

An analytic calculation of the growth index for $f(R)$ dark energy model

Qing-Guo Huang^a

State Key Laboratory of Theoretical Physics, Institute of Theoretical Physics, Chinese Academy of Science, Beijing 100190, People's Republic of China

Received: 29 May 2014 / Accepted: 2 July 2014 / Published online: 17 July 2014
© The Author(s) 2014. This article is published with open access at Springerlink.com

Abstract We derive the analytic formula of the growth index for the $f(R)$ dark energy model where the effect on the growth of matter density perturbation δ_m from modified gravity (MG) is encoded in the effective Newton coupling constant G_{eff} in MG (or equivalently $g \equiv G_{\text{eff}}/G$). Based on the analytic formula, we propose that the parameter g can be directly found by comparing the observed growth rate $f_g \equiv d \ln \delta_m / d \ln a$ to the prediction of f_g in general relativity.

1 Introduction

The accelerating expansion of the present Universe was discovered by studying type Ia supernovae [1, 2]. Up to now, the standard Λ CDM model in the framework of general relativity (GR) is able to explain the present cosmic acceleration within observational errors. However, how to explain the tiny value of the cosmological constant compared to the known physical scales is still a big challenge.

Modified gravity (MG), for example $f(R)$ gravity, provides a geometrical origin to the present cosmic acceleration. The basic idea of MG dark energy is that gravity is modified on the cosmological scales when the Ricci scalar R is of the order of today's Ricci scalar R_0 , while GR is recovered in the region of $R \gg R_0$. However, it is quite non-trivial to construct a viable $f(R)$ dark energy model which is consistent with both cosmological and local gravity constraints. See some typical viable $f(R)$ dark energy models in [3–9]. It is useful to introduce the effective equation of state parameter $w = p_{\text{de}}/\rho_{\text{de}}$ to describe the difference between the Friedmann–Robertson–Walker (FRW) background evolutions of MG and the standard Λ CDM model, where the effective pressure p_{de} and the energy density ρ_{de} are determined by using the Einsteinian representation of the gravita-

tional field equations. On the other hand, since the gravity in MG is different from GR, the evolution of the matter density perturbation $\delta_m \equiv \delta\rho_m/\rho_m$ provides a crucial tool to distinguish the MG dark energy model from the dark energy model in GR, in particular the standard Λ CDM model. For simplicity, the growth rate f_g of the matter density perturbation can be parametrized by [10]

$$f_g \equiv \frac{d \ln \delta_m}{d \ln a} \equiv \Omega_m(z)^{\gamma(z)}, \quad (1)$$

where a is the scale factor, $\Omega_m(z)$ is the density parameter for dust-like matter at redshift z , and $\gamma(z)$ is the so-called growth index. In the Λ CDM model in GR, $w = -1$, and we have [11, 12]

$$\gamma \simeq 6/11. \quad (2)$$

Generically the effect on the matter density perturbation in MG is encoded in the effective Newton coupling constant G_{eff} . For simplicity, we introduce a new quantity $g \equiv G_{\text{eff}}/G$ to measure the difference between MG and GR. In general, w is time dependent and g is time and scale dependent in MG, and then the growth index γ is expected to be time and scale dependent. During a deep matter-dominant era GR is recovered, while the gravity is modified in the low redshift era when the cosmic acceleration occurs. One can expect that the evolutions of both FRW background and matter density perturbation in MG are too complicated to be solved analytically from the deep matter-dominant era to the accelerating era.

In this paper we focus on the growth of matter density perturbation in the $f(R)$ dark energy model. We suppose that g is parametrized as follows:

$$g = g_0 + g_1(1 - \Omega_m), \quad (3)$$

where g_0 and g_1 are two constants. Here g_1 is used to characterize the time-evolution of g . Note that both $g_0 = g_0(k)$ and $g_1 = g_1(k)$ are scale dependent generically. In the deep

^a e-mail: huangqg@itp.ac.cn

matter-dominant era ($\Omega_m \rightarrow 1$), GR should be recovered and then $g \rightarrow 1$. But g can deviate from 1 at low redshift. This parameterization can cover many viable $f(R)$ dark energy models at low redshift. Based on such a parameterization, we analytically solve the equation of motion of δ_m and work out an analytic formula for the growth index. Furthermore, we find that g can be directly found by comparing the observed growth rate f_g to the prediction of f_g in GR.

This paper will be organized as follows. In Sect. 2 we briefly review the $f(R)$ dark energy model. In Sect. 3 we analytically calculate the growth index for $f(R)$ dark energy model. A summary and discussion are given in Sect. 4.

2 A brief introduction to the $f(R)$ dark energy model

Let us start with the following action:

$$S = \frac{1}{16\pi G} \int d^4x \sqrt{-g} f(R) + S_m, \tag{4}$$

where G is the Newton coupling constant, S_m is the action for the matter, $R = 6(2H^2 + \dot{H})$ and H denotes the Hubble parameter. If $f(R) = R - 2\Lambda$, the above action reduces to the Einstein–Hilbert action for the Λ CDM model in GR. In this paper we consider the case that $f(R)$ vanishes for $R = 0$, which implies that no cosmological constant is introduced. The $f(R)$ gravity contains a new scalar degree of freedom dubbed “scalaron” whose mass depends on the Ricci scalar R [13]. The stability of $f(R)$ theory requires

$$F \equiv f_{,R} > 0, \quad F_{,R} \equiv f_{,RR} > 0, \tag{5}$$

where $f_{,R} = df(R)/dR$ and $F_{,R} = dF(R)/dR$. The former condition implies that gravity is attractive and graviton is not a ghost, and the latter condition means that the scalaron is not a tachyon. In addition, the viable $f(R)$ dark energy model is required to be similar to the Λ CDM model during the radiation and deep matter-dominant era, but important observable deviations from the Λ CDM model appear at low redshift. In order to measure such a deviation, we can introduce a dimensionless quantity defined by $\beta \equiv RF_{,R}/F$ which satisfies $0 < \beta < 1$ [3, 14].

Considering that S_m describes the dust-like matter (the pressure of dust-like matter equals 0), the equations of motion for the FRW background take the form

$$H^2 = \frac{1}{3} \left[\frac{1}{2}(FR - f) - 3H\dot{F} - 3(F - 1)H^2 \right] + \frac{8\pi G}{3} \rho_m, \tag{6}$$

$$-2\dot{H} = \ddot{F} - H\dot{F} + 2(F - 1)\dot{H} + 8\pi G\rho_m. \tag{7}$$

Here we focus on the late time Universe where the radiation can be ignored. From these two equations, the effective

energy density and pressure of $f(R)$ dark energy are, respectively, given by

$$\rho_{de} = \frac{1}{8\pi G} \left[\frac{1}{2}(FR - f) - 3H\dot{F} - 3(F - 1)H^2 \right], \tag{8}$$

$$p_{de} = -\rho_{de} + \frac{1}{8\pi G} [\ddot{F} - H\dot{F} + 2(F - 1)\dot{H}], \tag{9}$$

and then the effective equation of state parameter w reads

$$w = -1 + \frac{\ddot{F} - H\dot{F} + 2(F - 1)\dot{H}}{\frac{1}{2}(FR - f) - 3H\dot{F} - 3(F - 1)H^2}. \tag{10}$$

Combining Eqs. (6) and (7), the Ricci scalar becomes

$$R = 3[1 - 3w(1 - \Omega_m)]H^2, \tag{11}$$

where

$$\Omega_m \equiv \frac{8\pi G\rho_m}{3H^2} \tag{12}$$

is the density parameter for the dust-like matter.

Many typical viable $f(R)$ dark energy models which are consistent with both cosmological and local gravity constraints are summarized in [14]. All of them can be written in the following form:

$$f(R) = R - \lambda R_s Y(x), \tag{13}$$

where $x = R/R_s$, $R_s (> 0)$ is a characteristic value of R and λ is a positive parameter. The function $Y(x)$ in the viable model takes the forms: (i) $Y(x) = x^p$ ($0 < p < 1$) [3], (ii) $Y(x) = x^{2n}/(x^{2n} + 1)$ ($n > 0$) [4], (iii) $Y(x) = 1 - (1 + x^2)^{-n}$ ($n > 0$) [6], (iv) $Y(x) = 1 - e^{-x}$ [8, 9], (v) $Y(x) = \tanh(x)$ [7], etc. We find that all of these models satisfy $F = f_{,R} < 1$.

3 The analytic formula of the growth index for $f(R)$ dark energy model

From now on, we will focus on the evolution of the matter density perturbation in MG. In the sub-horizon limit, the evolution of the matter density fluctuation, δ_m , is governed by

$$\ddot{\delta}_m + 2H\dot{\delta}_m - 4\pi G_{\text{eff}}\rho_m\delta_m = 0, \tag{14}$$

$$G_{\text{eff}} = g(a, k, R) \cdot G, \tag{15}$$

where

$$g(a, k, R) \equiv \frac{1}{F} \left(1 + \frac{1}{3} \frac{1}{1 + \frac{M^2 a^2}{k^2}} \right), \tag{16}$$

in [15, 16], or without taking any approximation at the matter-dominant stage [17, 18]

$$g(a, k, R) = 1 + \frac{8}{3} \frac{4/F - 3}{27 + 8 \left(\frac{k^2}{a^2 M^2} \right)^4} \left(\frac{k^2}{a^2 M^2} \right)^4, \tag{17}$$

and

$$M^2 = \frac{R}{3\beta} = \frac{F}{3F_{,R}}. \tag{18}$$

Here M is nothing but the mass of the scalaron. The positivity of both $f_{,R}$ and $f_{,RR}$ guarantees the positivity of the mass square of the scalaron. In the scales which are much smaller than M^{-1} , GR is recovered, and the gravity is modified in the scales around or larger than M^{-1} . In addition, considering (16) or (17) with $F < 1$, we find $g > 1$ for the viable $f(R)$ dark energy models in the literature.

In the deep matter-dominant era, $a \sim t^{2/3}$, and then Eq. (14) becomes

$$\ddot{\delta}_m + \frac{4}{3t}\dot{\delta}_m - \frac{2}{3t^2}g\delta_m = 0, \tag{19}$$

whose solution is given by $\delta_m \sim t^{\frac{\sqrt{1+24g}-1}{6}} \sim a^{\frac{\sqrt{1+24g}-1}{4}}$, where g is taken as a constant. In this era, GR is proposed to be recovered ($g \rightarrow 1$) and then $\delta_m \sim a$.

Now let us switch to the late time Universe where the energy densities of effective dark energy and dust-like matter are comparable to each other. Equation (14) can be re-written as follows:

$$\frac{d^2 \ln \delta_m}{d \ln a^2} + \left(\frac{d \ln \delta_m}{d \ln a} \right)^2 + \frac{d \ln \delta_m}{d \ln a} \left[\frac{1}{2} - \frac{3}{2}w(1 - \Omega_m) \right] = \frac{3}{2}g\Omega_m, \tag{20}$$

or equivalently

$$\frac{df_g}{d \ln a} + f_g^2 + f_g \left[\frac{1}{2} - \frac{3}{2}w(1 - \Omega_m) \right] = \frac{3}{2}g\Omega_m. \tag{21}$$

From Eqs. (6), (7), and (12), we have

$$\frac{d\Omega_m}{d \ln a} = 3w\Omega_m(1 - \Omega_m). \tag{22}$$

Combining with the definition of the growth index γ in Eqs. (1), (20) becomes

$$3w\Omega_m(1 - \Omega_m) \ln \Omega_m \frac{d\gamma}{d\Omega_m} + 3w \left(\gamma - \frac{1}{2} \right) (1 - \Omega_m) + \Omega_m^\gamma - \frac{3}{2}g\Omega_m^{1-\gamma} + \frac{1}{2} = 0. \tag{23}$$

Usually the form of $g(a, k, R)$ is expected to be very complicated. In order to capture the main feature of the $f(R)$ dark energy model at low redshift, we expand g as a power series about $(1 - \Omega_m) \sim 0$ for a given perturbation mode k ,

$$g = \sum_{n=0} g_n(1 - \Omega_m)^n. \tag{24}$$

In this paper we take the first two terms like that in Eq. (3) into account.¹ For a slowly varying equation of the state parameter

¹ The case with $g_0 = 1$ is discussed in [19].

w ($|dw/d\Omega_m| \ll (1 - \Omega_m)$), the solution of Eq. (23) takes the form

$$\gamma = \frac{c_{-1}}{1 - \Omega_m} + c_0 + c_1(1 - \Omega_m) + \mathcal{O}\left((1 - \Omega_m)^2\right), \tag{25}$$

where c_{-1} , c_0 , and c_1 can be calculated order by order,

$$c_{-1} = \ln \frac{1 + \sqrt{1 + 24g_0}}{6g_0}, \tag{26}$$

$$c_0 = 1 - \frac{c_{-1}}{2} + \frac{-2 + 3e^{c_{-1}w} - 3e^{2c_{-1}g_1}}{2 + 3e^{2c_{-1}g_0} - 6e^{c_{-1}w}}, \tag{27}$$

$$c_1 = \frac{(c_{-1}^2 + 4c_{-1}c_0 - 4(1 - c_0)c_0)(2 - 3e^{2c_{-1}g_0})}{8(2 + 3e^{2c_{-1}g_0} - 12e^{c_{-1}w})} - \frac{c_{-1}(4 - 3e^{2c_{-1}g_0} - 6e^{c_{-1}w})}{6(2 + 3e^{2c_{-1}g_0} - 12e^{c_{-1}w})} + \frac{3g_1(1 - c_{-1}/2 - c_0)e^{2c_{-1}}}{2 + 3e^{2c_{-1}g_0} - 12e^{c_{-1}w}}. \tag{28}$$

The expressions of the higher order terms are quite complicated, and the readers can easily work them out once they need to. Here the first two terms on the right hand side of Eq. (25) make the main contributions to γ and the term with c_1 is roughly negligible if both $(g_0 - 1)$ and g_1 are much less than 1. If $g_0 \neq 1$, the growth index is expected to be time-evolving, and the ansatz with a constant growth index is not generic for $f(R)$ dark energy model. Our analytic formula indicates that a better ansatz for γ is

$$\gamma(z) \simeq \frac{\gamma_{-1}}{1 - \Omega_m} + \gamma_0 + \gamma_1(1 - \Omega_m), \tag{29}$$

where γ_{-1} , γ_0 and γ_1 are constants.

For $g_0 = 1$ and $g_1 = 0$, our result reduces to GR where $c_{-1} = 0$, $c_0 = \frac{3(1-w)}{5-6w}$ and $c_1 = \frac{3}{125} \frac{(1-w)(1-3w/2)}{(1-6w/5)^2(1-12w/5)}$.² For $g_0 = 1$,

$$c_{-1} = 0, \quad c_0 = \frac{3(1 - g_1 - w)}{5 - 6w}. \tag{30}$$

In the Dvali–Gabadadze–Porrati (DGP) model [20,21], $g = 1 - \frac{1}{3} \frac{1 - \Omega_m^2}{1 + \Omega_m^2}$. In the matter-dominant era, $g \rightarrow 1 - \frac{1}{3}(1 - \Omega_m)$ and $w \rightarrow -1/2$, and thus $\gamma \simeq 11/16$ which is the same as that in [10].

Nowadays the property of dark energy has been tightly constrained from observations [22]. The Λ CDM model can fit the data, and the room for $f(R)$ dark energy model has been tightly constrained, for example $|g_0 - 1| \lesssim \mathcal{O}(0.1)$ and $|g_1| \lesssim \mathcal{O}(0.1)$. Therefore c_{-1} and c_0 can be expanded around the case of GR ($g_0 = 1$ and $g_1 = 0$),

² The denominator of c_1 is slightly different from that in [11] where $c_1 = \frac{3}{125} \frac{(1-w)(1-3w/2)}{(1-6w/5)^3}$.

$$c_{-1} = -\frac{3}{5}(g_0 - 1) + \mathcal{O}((g_0 - 1)^2), \tag{31}$$

$$c_0 = \frac{3(1 - g_1 - w)}{5(1 - 6w/5)} + \frac{9(7 + 18g_1 - 20w - 12g_1w + 12w^2)}{250(1 - 6w/5)^2}(g_0 - 1) + \mathcal{O}((g_0 - 1)^2). \tag{32}$$

For $w = -1$ and $g_1 = 0$, $c_0 \simeq \frac{6}{11} + \frac{351}{1,210}(g_0 - 1)$.

Applying our analytic formula in (25) to Eq. (1), we can easily calculate the growth rate $f_g(z)$. Testing the growth rate f_g from our analytic result in (25) against the value obtained by numerical calculation, the accuracy at low redshift ($z \lesssim 1$) is better than a few percents. Combining Eq. (25) with Eq. (1) and expanding f_g up to the order of $(1 - \Omega_m)^2$, we have

$$f_g(z) = e^{-c_{-1} - (c_{-1}/2 + c_0)(1 - \Omega_m) + \mathcal{O}((1 - \Omega_m)^2)}, \tag{33}$$

where

$$\frac{c_{-1}}{2} + c_0 = \frac{3(1 - g_1 - w)}{5(1 - 6w/5)} - \frac{3(2 - 27g_1 + 18g_1w)}{125(1 - 6w/5)^2}(g_0 - 1) + \mathcal{O}((g_0 - 1)^2). \tag{34}$$

Since the first term on the right hand side of Eq. (34) is dominant, we have

$$\frac{f_{g, \text{MG}}(z)}{f_{g, \text{GR}}(z)} \simeq \exp\left[-c_{-1} + \frac{3g_1\Omega_{\text{de}}(z)}{5 - 6w}\right], \tag{35}$$

where $\Omega_{\text{de}} = (1 - \Omega_m)$ is the dark energy density parameter. For $g_1 = 0$ and $g_0 > 1$ (or equivalently $c_{-1} < 0$), which implies that gravity is stronger than GR, the growth rate of the matter density perturbation is enhanced by a factor $e^{-c_{-1}}$. Motivated by Eq. (35), we propose

$$1 - \frac{5}{3}c_{-1} + \frac{g_1\Omega_{\text{de}}(z)}{1 - 6w/5} \simeq r_{g, \text{obs}}(z), \tag{36}$$

where

$$r_{g, \text{obs}}(z) \equiv 1 + \frac{5}{3} \ln \frac{f_{g, \text{obs}}(z)}{f_{g, \text{GR}}(z)}. \tag{37}$$

Note that c_{-1} is a function of g_0 . Once we can construct the relation between $r_g(z)$ and Ω_{de} from cosmological observations, we can easily find g_0 and g_1 . Roughly speaking, the value of g_0 can be determined by the value of $r_{g, \text{obs}}(z)$ when $\Omega_{\text{de}} \simeq 0$, and g_1 is related to the tilt of $r_{g, \text{obs}}(z)$ at low redshift. If $|g_0 - 1| \lesssim \mathcal{O}(0.1)$, Eq. (36) becomes

$$g_0 + \frac{g_1\Omega_{\text{de}}(z)}{1 - 6w/5} \simeq r_{g, \text{obs}}(z), \tag{38}$$

which indicates that the redshift-independent part of $r_{g, \text{obs}}(z)$ is equal to g_0 .

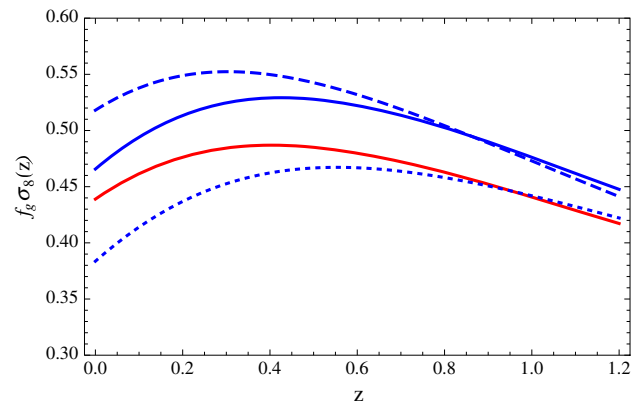


Fig. 1 The plot of $f_g \sigma_8(z)$. Here we adopt $w = -1$, $\Omega_m^0 = 0.315$ and $\sigma_8 = 0.829$. The red solid, blue dashed, blue dotted, and blue solid curves correspond to Λ CDM model, $g = 1.3$, $g = 1 - 0.5(1 - \Omega_m)$ and $g = 1.3 - 0.5(1 - \Omega_m)$ respectively

In the literature, one may prefer to constrain $f_g \sigma_8(z)$ from cosmological observations, where σ_8 is today's root-mean-square mass fluctuation on $8 h^{-1}$ Mpc. Because the equation of motion of δ_m is a linear equation, one can define a normalized growth function $D(z)$ via

$$D(z) \equiv \delta_m(z)/\delta_m(z = 0), \tag{39}$$

and then $\sigma_8(z) = \sigma_8 D(z)$. Solving Eq. (1), we get

$$D(z) = \frac{1}{1 + z} \exp\left[\int_0^z \left(1 - \Omega_m(z')\right)^{\gamma(z')} \frac{dz'}{1 + z'}\right]. \tag{40}$$

Therefore

$$f_g \sigma_8(z) = D(z) \Omega_m(z)^{\gamma(z)} \sigma_8. \tag{41}$$

Using our analytic result, $f_g \sigma_8(z)$ is plotted in Fig. 1.

Roughly speaking, if $|g_0 - 1| \lesssim 1$ and $|g_1| \lesssim 1$, g_0 shifts the amplitude of $f_g \sigma_8(z)$ and g_1 changes the shape of $f_g \sigma_8(z)$.

4 Summary and discussion

To summarize, we analytically calculate the growth index in the $f(R)$ dark energy model. Actually our results are applicable for more general MG dark energy models, for example $f(T)$ dark energy model [23–26], as long as the effect on the growth of matter density perturbation from MG is encoded in $g = G_{\text{eff}}/G$. As we know, there are two key parameters for MG dark energy model, namely w and G_{eff} (or equivalently g). The former parameter determines the expansion history of our Universe, and the latter parameter tells us how the matter density perturbation grows up. Adopting the analytic formula, we find a simple relation between g and the growth rate in Eq. (35), and then we propose that g can be

directly found by comparing the observed growth rate f_g to the prediction of f_g in GR. In the literature, one would like to use $f_g\sigma_8(z)$ to characterize the growth of the matter density perturbation. In this case one can also use our analytic formula to calculate $f_g\sigma_8(z)$ and then fit g_0 and A from the data.

Recently the anisotropic clustering of the Baryon Oscillation Spectroscopic Survey (BOSS) CMASS Data Release 11 (DR11) sample was analyzed. The combination of Planck and CMASS implies $\gamma = 0.772^{+0.124}_{-0.097}$ and a similar result $\gamma = 0.76 \pm 0.11$ is obtained when replacing Planck with WMAP9 in [27]. Both results deviate from the prediction of Λ CDM in GR at more than 2σ level. The large value of γ may come from the large value of σ_8 from Planck, or it is just a statistical fluctuation. Considering $f_g\sigma_8(z = 0.57) = 0.419 \pm 0.044$ from BOSS CMASS DR11, we obtain $g_0 \simeq 0.73$ in the reference Λ CDM model ($\Omega_m^0 = 0.315$ and $\sigma_8 = 0.829$) from Planck [22]. A careful data fitting will be done in the near future [28]. In a word, if such a deviation is confirmed in the future, we really need to modify the gravity.

Finally for some other aspects on $f(R)$ dark energy model see [29–44] etc.

Acknowledgments This work was initiated during High1-2014 KIAS-NCTS joint workshop on particle physics, string theory and cosmology. Q.-G.H. is supported by the project of Knowledge Innovation Program of Chinese Academy of Science and grants from NSFC (Grant No. 10821504, 11322545 and 11335012).

Open Access This article is distributed under the terms of the Creative Commons Attribution License which permits any use, distribution, and reproduction in any medium, provided the original author(s) and the source are credited.

Funded by SCOAP³ / License Version CC BY 4.0.

References

1. S. Perlmutter et al., Supernova cosmology project collaboration. *Astrophys. J.* **517**, 565 (1999). [astro-ph/9812133](#)
2. A.G. Riess et al., Supernova search team collaboration. *Astron. J.* **116**, 1009 (1998). [astro-ph/9805201](#)
3. L. Amendola, R. Gannouji, D. Polarski, S. Tsujikawa, *Phys. Rev. D* **75**, 083504 (2007). [gr-qc/0612180](#)
4. W. Hu, I. Sawicki, *Phys. Rev. D* **76**, 064004 (2007). [arXiv:0705.1158](#) [astro-ph]
5. S.A. Appleby, R.A. Battye, *Phys. Lett. B* **654**, 7 (2007). [arXiv:0705.3199](#) [astro-ph]
6. A.A. Starobinsky, *JETP Lett.* **86**, 157 (2007). [arXiv:0706.2041](#) [astro-ph]
7. S. Tsujikawa, *Phys. Rev. D* **77**, 023507 (2008). [arXiv:0709.1391](#) [astro-ph]
8. G. Cognola, E. Elizalde, S. Nojiri, S.D. Odintsov, L. Sebastiani, S. Zerbini, *Phys. Rev. D* **77**, 046009 (2008). [arXiv:0712.4017](#) [hep-th]
9. E.V. Linder, *Phys. Rev. D* **80**, 123528 (2009). [arXiv:0905.2962](#) [astro-ph.CO]
10. E.V. Linder, R.N. Cahn, *Astropart. Phys.* **28**, 481 (2007). [astro-ph/0701317](#)
11. L.-M. Wang, P.J. Steinhardt, *Astrophys. J.* **508**, 483 (1998). [astro-ph/9804015](#)
12. E.V. Linder, *Phys. Rev. D* **72**, 043529 (2005). [astro-ph/0507263](#)
13. A.A. Starobinsky, *Phys. Lett. B* **91**, 99 (1980)
14. S. Tsujikawa, R. Gannouji, B. Moraes, D. Polarski, *Phys. Rev. D* **80**, 084044 (2009). [arXiv:0908.2669](#) [astro-ph.CO]
15. P. Zhang, *Phys. Rev. D* **73**, 123504 (2006). [astro-ph/0511218](#)
16. S. Tsujikawa, *Phys. Rev. D* **76**, 023514 (2007). [arXiv:0705.1032](#) [astro-ph]
17. A. de la Cruz-Dombriz, A. Dobado, A.L. Maroto, *Phys. Rev. D* **77**, 123515 (2008). [arXiv:0802.2999](#) [astro-ph]
18. H. Motohashi, A.A. Starobinsky, J.'i. Yokoyama, *Int. J. Mod. Phys. D* **18**, 1731 (2009). [arXiv:0905.0730](#) [astro-ph.CO]
19. P.G. Ferreira, C. Skordis, *Phys. Rev. D* **81**, 104020 (2010). [arXiv:1003.4231](#) [astro-ph.CO]
20. G.R. Dvali, G. Gabadadze, M. Porrati, *Phys. Lett. B* **485**, 208 (2000). [hep-th/0005016](#)
21. C. Deffayet, G.R. Dvali, G. Gabadadze, *Phys. Rev. D* **65**, 044023 (2002). [astro-ph/0105068](#)
22. P.A.R. Ade et al., [Planck Collaboration]. [arXiv:1303.5076](#) [astro-ph.CO]
23. E.V. Linder, *Phys. Rev. D* **81**, 127301 (2010) [Erratum-ibid. *D* **82**, 109902 (2010)]. [arXiv:1005.3039](#) [astro-ph.CO]
24. K. Bamba, C.-Q. Geng, C.-C. Lee, *JCAP* **1008**, 021 (2010). [arXiv:1005.4574](#) [astro-ph.CO]
25. R. Zheng, Q.-G. Huang, *JCAP* **1103**, 002 (2011). [arXiv:1010.3512](#) [gr-qc]
26. W.-S. Zhang, C. Cheng, Q.-G. Huang, M. Li, S. Li, X.-D. Li, S. Wang, *Sci. China Phys. Mech. Astron.* **55**, 2244 (2012). [arXiv:1202.0892](#) [astro-ph.CO]
27. F. Beutler et al. [BOSS Collaboration]. [arXiv:1312.4611](#) [astro-ph.CO]
28. Q.-G. Huang et al. (to appear)
29. A. De Felice, S. Tsujikawa, *Living Rev. Rel.* **13**, 3 (2010). [arXiv:1002.4928](#) [gr-qc]
30. S.'i. Nojiri, S.D. Odintsov, *Phys. Rept.* **505**, 59 (2011). [arXiv:1011.0544](#) [gr-qc]
31. D. Polarski, R. Gannouji, *Phys. Lett. B* **660**, 439 (2008). [arXiv:0710.1510](#) [astro-ph]
32. H. Wei, *Phys. Lett. B* **664**, 1 (2008). [arXiv:0802.4122](#) [astro-ph]
33. Y. Gong, *Phys. Rev. D* **78**, 123010 (2008). [arXiv:0808.1316](#) [astro-ph]
34. R. Gannouji, B. Moraes, D. Polarski, *JCAP* **0902**, 034 (2009). [arXiv:0809.3374](#) [astro-ph]
35. K. Bamba, C.-Q. Geng, S.'i. Nojiri, S.D. Odintsov, *Phys. Rev. D* **79**, 083014 (2009). [arXiv:0810.4296](#) [hep-th]
36. P. Wu, H.W. Yu, X. Fu, *JCAP* **0906**, 019 (2009). [arXiv:0905.3444](#) [gr-qc]
37. T. Biswas, E. Gerwick, T. Koivisto, A. Mazumdar, *Phys. Rev. Lett.* **108**, 031101 (2012). [arXiv:1110.5249](#) [gr-qc]
38. K. Bamba, A. Lopez-Revelles, R. Myrzakulov, S.D. Odintsov, L. Sebastiani, *Class. Quant. Grav.* **30**, 015008 (2013). [arXiv:1207.1009](#) [gr-qc]
39. J.-h. He, B. Wang, *Phys. Rev. D* **87**, 023508 (2013). [arXiv:1208.1388](#) [astro-ph.CO]
40. S. Basilakos, S. Nesseris, L. Perivolaropoulos, *Phys. Rev. D* **87**(12), 123529 (2013). [arXiv:1302.6051](#) [astro-ph.CO]
41. A. Abebe, A. de la Cruz-Dombriz, P.K.S. Dunsby, *Phys. Rev. D* **88**, 044050 (2013). [arXiv:1304.3462](#) [astro-ph.CO]
42. L. Xu, [arXiv:1306.2683](#) [astro-ph.CO]
43. J. Dossett, B. Hu, D. Parkinson, [arXiv:1401.3980](#) [astro-ph.CO]
44. A. Pouri, S. Basilakos, M. Plionis, [arXiv:1402.0964](#) [astro-ph.CO]