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Thermodynamics of Lovelock black holes with a nonminimal scalar field

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ABSTRACT: We source the Lovelock gravity theories indexed by an integer k and fixed by requiring a unique anti-de Sitter vacuum with a self-interacting nonminimal scalar field in arbitrary dimension d . For each inequivalent Lovelock gravity theory indexed by the integer k , we establish the existence of a two-parametric self-interacting potential that permits to derive a class of black hole solutions with planar horizon for any arbitrary value of the nonminimal coupling parameter. In the thermodynamical analysis of the solution, we show that, once regularized the Euclidean action, the mass contribution coming from the gravity side exactly cancels, order by order, the one arising from the matter part yielding to a vanishing mass. This result is in accordance with the fact that the entropy of the solution, being proportional to the lapse function evaluated at the horizon, also vanishes. Consequently, the integration constant appearing in the solution is interpreted as a sort of hair which turns out to vanish at high temperature.

KEYWORDS: Classical Theories of Gravity, Black Holes

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1 Introduction

During the last decades, the interests on higher-dimensional physics have grown up particularly concerning the higher-dimensional General Relativity. In this context, apart from the standard Einstein-Hilbert action, there also exists an interesting gravity theory in dimension $d \geq 5$ involving higher powers of the curvatures such that the field equations for the metric are at most of second-order. This theory known as the Lovelock gravity has been first implemented in five dimensions by Lanczos [1] and then generalized in higher dimension $d \geq 5$ by Lovelock [2]. The resulting action is a d -form constructed out of the vielbein, the spin connection and their exterior derivative without using the Hodge dual. The invariance under local Lorentz transformations of the Lovelock Lagrangian can be extended to a local anti-de Sitter (AdS) or Poincaré symmetry through a particular choice of the coefficients appearing in the Lovelock expansion. In both cases, the resulting Lagrangian is a Chern-Simons form whose supersymmetric extensions are also known; see [3] for a good review on Chern-Simons (super)gravity. The Lovelock gravity or its Chern-Simons particular case have been shown to possess (topological) AdS black hole solutions with interesting thermodynamical properties [4–7] generalizing those obtained in the Einstein-Gauss-Bonnet case [8, 9].

In order to source the Lovelock gravity with the purpose of obtaining black hole solutions, nonminimally coupled scalar field can be an excellent candidate. Indeed, the nonminimal coupling may be useful to escape the standard no-hair theorems, and black hole solutions are known in this case in standard Einstein gravity (with and without cosmological constant) in four dimensions with a conformal coupling [10–13]. Recently, there has been a renewal interest concerning this kind of source particularly in order to gain a better understanding of some unconventional superconductors [14–16]. Indeed, it is believed that black hole solutions with scalar hair at low temperature that disappears at high temperature will be of particular importance for the unconventional superconductors since they will reproduce the correct behavior of the superconductor phase diagram.

In the present work, we consider a series of particular Lovelock gravity actions indexed by an integer k in dimension d , and sourced by a self-interacting nonminimal scalar field. To be more precise, the arbitrary coefficients appearing in the Lovelock series are fixed by requiring that the resulting theory has a unique anti-de Sitter vacuum with a fixed cosmological constant. This yields to $k = 1, 2 \dots, [\frac{d-1}{2}]$ inequivalent gravity theories where $k = 1$ corresponds to the standard Einstein-Hilbert action. Note that such model with a mass term potential has been considered in [17, 18], and black hole solutions with planar horizon have been obtained for two particular values of the nonminimal coupling parameter. Here, we generalize these solutions and obtain black hole solutions for any value of the nonminimal coupling parameter. This task is achieved thanks to the introduction of a much more general self-interacting term that depends explicitly on the integer index k . More specifically, we show that for each inequivalent Lovelock gravity theory indexed by k , there exists a judicious choice for the self-interacting potential that permits to derive black hole solutions for arbitrary value of the nonminimal coupling parameter. The solutions are shown to be uniparametric and reduce to those derived in [17, 18] for the special values of the nonminimal coupling parameter. The thermodynamics study of the solutions realized through the Hamiltonian analysis reveals that these configurations have a zero mass as well as a vanishing entropy. More precisely, we show that the mass contribution coming from the gravity side exactly cancels the mass contribution inherent to the matter source. The vanishing of the entropy can be also corroborated by the fact that the entropy of the solutions is proportional to the lapse function evaluated at the horizon. Since the constant appearing in the black hole solution does not contribute to any conserved charge, we interpret it as a sort of hair which turns out to be inversely proportional to the temperature. Hence, at high temperature, this hair will disappear and this kind of behavior is expected in the unconventional superconductors in order to correctly reproduce the phase diagram.

The plan of the paper is organized as follows. In the next section, we present the model, the field equations and the solutions. Section 3 is devoted to the thermodynamical analysis while our conclusions and comments are given in the last section.

2 Lovelock black hole solution with arbitrary nonminimal coupling parameter

We consider a generalization of the Einstein-Hilbert gravity action in arbitrary dimension d yielding at most to second-order field equations for the metric and known as the Lovelock Lagrangian. This latter is a d -form constructed with the vielbein e^a , the spin connection ω^{ab} , and their exterior derivatives without using the Hodge dual. The Lovelock action is a polynomial of degree $[d/2]$ (where $[x]$ denotes the integer part of x) in the curvature two-form, $R^{ab} = d\omega^{ab} + \omega_c^a \wedge \omega^{cb}$ as

$$\int \sum_{p=0}^{[d/2]} \alpha_p L^{(p)}, \tag{2.1a}$$

$$L^{(p)} = \epsilon_{a_1 \dots a_d} R^{a_1 a_2} \dots R^{a_{2p-1} a_{2p}} e^{a_{2p+1}} \dots e^{a_d}, \tag{2.1b}$$

where the α_p are arbitrary dimensionful coupling constants and where wedge products between forms are understood. Here $L^{(0)}$ and $L^{(1)}$ are proportional respectively to the cosmological term and the Einstein-Hilbert Lagrangian. As shown in ref. [6], requiring the Lovelock action to have a unique AdS vacuum with a unique cosmological constant, fixes the α_p yielding to a series of actions indexed by an integer k , and given by

$$I_k = -\frac{1}{2k(d-3)!} \int \sum_{p=0}^k \frac{C_p^k}{(d-2p)} L^{(p)}, \quad 1 \leq k \leq \left\lfloor \frac{d-1}{2} \right\rfloor, \quad (2.2)$$

where C_p^k corresponds to the combinatorial factor. The global factor in front of the integral is chosen such that the gravity action (2.2) can be re-written in the standard fashion as

$$I_k = \frac{1}{2} \int d^d x \sqrt{-g} \left[R + \frac{(d-1)(d-2)}{k} + \frac{(k-1)}{2(d-3)(d-4)} \mathcal{L}_{GB} + \frac{(k-1)(k-2)}{3!(d-3)(d-4)(d-5)(d-6)} \mathcal{L}_{(3)} + \dots \right], \quad (2.3)$$

where $\mathcal{L}_{GB} = R^2 - 4R_{\mu\nu}R^{\mu\nu} + R_{\alpha\beta\mu\nu}R^{\alpha\beta\mu\nu}$ stands for the Gauss-Bonnet Lagrangian, and $\mathcal{L}_{(3)}$ is given by

$$\begin{aligned} \mathcal{L}_{(3)} = & R^3 - 12RR_{\mu\nu}R^{\mu\nu} + 16R_{\mu\nu}R^\mu{}_\rho R^{\nu\rho} + 24R_{\mu\nu}R_{\rho\sigma}R^{\mu\rho\nu\sigma} + 3RR_{\mu\nu\rho\sigma}R^{\mu\nu\rho\sigma} \\ & - 24R_{\mu\nu}R^\mu{}_{\rho\sigma\kappa}R^{\nu\rho\sigma\kappa} + 4R_{\mu\nu\rho\sigma}R^{\mu\nu\eta\zeta}R^{\rho\sigma}{}_{\eta\zeta} - 8R_{\mu\rho\nu\sigma}R^\mu{}_\nu{}^\nu{}_\zeta R^{\rho\eta\sigma\zeta}. \end{aligned}$$

Note that in odd dimension $d = 2n - 1$ and for $k = n - 1$, the corresponding action I_{n-1} is a Chern-Simons action, that is a $(2n - 1)$ -form whose exterior derivative can be written as the contraction of an invariant tensor with the wedge product of n curvatures two-forms. In even dimension $d = 2n$, the maximal value of k is $n - 1$, and in this case the resulting gravity action has a Born-Infeld like structure since it can be written as the Pfaffian of the 2-form $\bar{R}^{ab} = R^{ab} + e^a e^b$. The gravity theories I_k have been shown to possess black hole solutions with interesting features, in particular concerning their thermodynamics properties, see [6] and [7].

For each $k \geq 2$, we source the gravity actions I_k with a self-interacting and nonminimally coupled scalar field, that is

$$S_k = I_k - \int d^d x \sqrt{-g} \left[\frac{1}{2} \partial_\mu \Phi \partial^\mu \Phi + \frac{\xi}{2} R \Phi^2 + U_k(\Phi) \right], \quad (2.4)$$

where ξ denotes the nonminimal coupling parameter and U_k is a potential term which depends explicitly on the index k , and whose form will be given below. The field equations read

$$\mathcal{G}_{\mu\nu}^{(k)} = T_{\mu\nu}, \quad (2.5a)$$

$$\square \Phi = \xi R \Phi + \frac{dU_k}{d\Phi}, \quad (2.5b)$$

where $\mathcal{G}_{\mu\nu}^{(k)}$ is the gravity tensor associated to the variation of the action I_k (2.2),

$$\begin{aligned} \mathcal{G}_{\mu\nu}^{(k)} &= G_{\mu\nu} - \frac{(d-1)(d-2)}{2k} g_{\mu\nu} + \frac{(k-1)}{2(d-3)(d-4)} K_{\mu\nu} \\ &+ \frac{(k-1)(k-2)}{3!(d-3)(d-4)(d-5)(d-6)} S_{\mu\nu} + \dots \end{aligned}$$

where $K_{\mu\nu}$ is the Gauss-Bonnet tensor

$$K_{\mu\nu} = 2(RR_{\mu\nu} - 2R_{\mu\rho}R^\rho{}_\nu - 2R^{\rho\sigma}R_{\mu\rho\nu\sigma} + R_\mu{}^{\rho\sigma\gamma}R_{\nu\rho\sigma\gamma}) - \frac{1}{2}g_{\mu\nu}\mathcal{L}_{GB}$$

and $S_{\mu\nu}$ arises from the variation of $\mathcal{L}_{(3)}$,

$$\begin{aligned} S_{\mu\nu} &= 3\left(R^2R_{\mu\nu} - 4RR_{\rho\mu}R^\rho{}_\nu - 4R^{\rho\sigma}R_{\rho\sigma}R_{\mu\nu} + 8R^{\rho\sigma}R_{\rho\mu}R_{\sigma\nu} - 4RR^{\rho\sigma}R_{\rho\mu\sigma\nu}\right. \\ &+ 8R^{\rho\kappa}R^\sigma{}_\kappa R_{\rho\mu\sigma\nu} - 16R^{\rho\sigma}R^\kappa{}_{(\mu}R_{|\kappa\sigma\rho|\nu)} + 2RR^{\rho\sigma\kappa}{}_\mu R_{\rho\sigma\kappa\nu} + R_{\mu\nu}R^{\rho\sigma\kappa\eta}R_{\rho\sigma\kappa\eta} \\ &- 8R^\rho{}_{(\mu}R^{\sigma\kappa\eta}{}_{|\rho|}R_{|\sigma\kappa\eta|\nu)} - 4R^{\rho\sigma}R^{\kappa\eta}{}_{\rho\mu}R_{\kappa\eta\sigma\nu} + 8R_{\rho\sigma}R^{\rho\kappa\sigma\eta}R_{\kappa\mu\eta\nu} - 8R_{\rho\sigma}R^{\rho\kappa\eta}{}_\mu R^\sigma{}_{\kappa\eta\nu} \\ &\left. + 4R^{\rho\sigma\kappa\eta}R_{\rho\sigma\zeta\mu}R_{\kappa\eta}{}^\zeta{}_\nu - 8R^{\rho\kappa\sigma\eta}R^\zeta{}_{\rho\sigma\mu}R_{\zeta\kappa\eta\nu} - 4R^{\rho\sigma\kappa}{}_\eta R_{\rho\sigma\kappa\zeta}R^\eta{}_\mu{}^\zeta{}_\nu\right) - \frac{1}{2}g_{\mu\nu}\mathcal{L}_{(3)}. \end{aligned}$$

In the matter part of the equations, $T_{\mu\nu}$ stands for the energy-momentum tensor of the scalar field whose expression is given by

$$T_{\mu\nu} = \partial_\mu\Phi\partial_\nu\Phi - g_{\mu\nu}\left(\frac{1}{2}\partial_\sigma\Phi\partial^\sigma\Phi + U_k(\Phi)\right) + \xi(g_{\mu\nu}\square - \nabla_\mu\nabla_\nu + G_{\mu\nu})\Phi^2. \quad (2.6)$$

For each inequivalent theory $k \geq 2$ and for $\xi \neq \frac{1}{4}$, the potential is given by the sum of the six following terms

$$U_k(\Phi) = \frac{1}{(4\xi-1)^2} \left[\alpha_1\Phi^2 + \alpha_2 b\Phi^{\frac{1}{2\xi}} + \alpha_3 b^2\Phi^{\frac{1-2\xi}{\xi}} + \alpha_4\Phi^{\frac{2k}{k-1}} + \alpha_5 b\Phi^{\frac{4\xi+k-1}{2\xi(k-1)}} + \alpha_6 b^2\Phi^{\frac{2\xi(2-k)+k-1}{\xi(k-1)}} \right] \quad (2.7)$$

and depends of a parameter b and the constants α_i which read

$$\begin{aligned} \alpha_1 &= 8\xi d(d-1)(\xi - \xi_d)(\xi - \xi_{d+1}), & \alpha_2 &= -16\xi^2(d-1)(\xi - \xi_d), & \alpha_3 &= 2\xi^2 \\ \alpha_4 &= -\frac{8\xi^{\frac{k}{k-1}}(1 + (k-1)d)((k-1)d + 2 - k)}{(k-1)k} (\xi - \hat{\xi}_{k,d}) (\xi - \hat{\xi}_{k,d+1}) \\ \alpha_5 &= \frac{16\xi(k(d-1) - (d-2)) - 4d(k-1) + 8(k-1)}{(k-1)} \xi^{\frac{2k-1}{k-1}}, & \alpha_6 &= -\frac{8\xi + 2(k-1)}{k-1} \xi^{\frac{2k-1}{k-1}}. \end{aligned}$$

Here ξ_d denotes the conformal coupling in d dimensions and we have defined $\hat{\xi}_{k,d}$ by

$$\xi_d = \frac{d-2}{4(d-1)}, \quad \hat{\xi}_{k,d} = \frac{(d-2)(k-1)}{4[(d-1)k - (d-2)]}. \quad (2.8)$$

Various comments can be made concerning the particular form of the potential (2.7). We anticipate that this potential naturally emerges looking for solutions of the field equations with an Ansatz of the form

$$ds^2 = -F(r)dt^2 + \frac{dr^2}{F(r)} + r^2 d\vec{x}_{d-2}^2, \quad \Phi = \Phi(r). \quad (2.9)$$

We also note that the self-interacting term depends explicitly on a coupling constant b as well as on the index parameter of the Lovelock theories k , and this dependence on k only concerns the last three terms of the potential. The first three terms of U_k exactly reproduce the potential that appears in many different contexts involving nonminimally coupled scalar field as for examples for the stealth configuration [19] on the BTZ black hole [20], or when looking for AdS wave backgrounds [21], or for black hole solutions when adding axionic fields coupled with the scalar field [22]. For $b = 0$ and $\xi = \hat{\xi}_{k,d}$ or $\xi = \hat{\xi}_{k,d+1}$, the potential reduces to the mass term considered in [17, 18] for which black hole solutions have been found. It is also clear from its expression that U_k is not well defined for $\xi = \frac{1}{4}$, and its derivation for this particular coupling involves logarithmic pieces.

Let us now present some solutions of the field equations (2.5). For $k \geq 2$ and $\xi \neq \frac{1}{4}$, an uniparametric black hole solution with planar horizon is given by

$$ds^2 = -r^2 \left(1 - \left[\xi(ar + b)^{\frac{4\xi}{4\xi-1}} \right]^{\frac{1}{k-1}} \right) dt^2 + \frac{dr^2}{r^2 \left(1 - \left[\xi(ar + b)^{\frac{4\xi}{4\xi-1}} \right]^{\frac{1}{k-1}} \right)} + r^2 d\vec{x}_{d-2}^2,$$

$$\Phi(r) = (ar + b)^{\frac{2\xi}{4\xi-1}}, \tag{2.10}$$

where a is an integration constant. We first stress that this solution is valid for any arbitrary value of the nonminimal coupling parameter $\xi \neq 0$, and for $\xi \in]0, \frac{1}{4}[$, the solution is asymptotically AdS while for $\xi > \frac{1}{4}$ the asymptotic behavior of the metric is faster than the usual AdS one. It is interesting to note that the singularity r_s is localized at $r_s = -\frac{b}{a}$ but it can always be hidden by the horizon $r_h = -\frac{b}{a} + \frac{\xi^{\frac{1-4\xi}{4\xi}}}{a}$ provided that the constant $a > 0$. For $b = 0$, and for $\xi = \hat{\xi}_{k,d}$ or for $\xi = \hat{\xi}_{k,d+1}$, the solutions (2.10) reduce to those found in the case of a mass term potential [17, 18]. We would like also to point out an important remark concerning the lapse function appearing in the metric solution that will be of importance in the thermodynamics study of these solutions. The metric function solution can be expressed in term of the scalar field as

$$r^2 \left(1 - \left[\xi(ar + b)^{\frac{4\xi}{4\xi-1}} \right]^{\frac{1}{k-1}} \right) = r^2 \left(1 - (\xi\Phi^2)^{\frac{1}{k-1}} \right)$$

and, hence the localization of the horizon r_h is such that

$$(1 - \xi\Phi^2)|_{r_h} = 0. \tag{2.11}$$

However, as we will see below, the entropy of the solutions is always proportional to this quantity (2.11), and consequently the solutions (2.10) will have a zero entropy. We also remark that the existence of these solutions is strongly inherent to the presence of the higher curvature terms present in the Lovelock Lagrangian. Indeed, it is clear from the different expressions obtained here that the standard GR limit $k = 1$ is singular, and hence these solutions are only effective for the higher order Lovelock terms $k \geq 2$.

To conclude this section, we present the solution for the coupling $\xi = \frac{1}{4}$. The self-interacting potential is given by

$$U_{[k, \xi = \frac{1}{4}]}(\Phi) = \frac{1}{8k(k-1)} \left\{ 4 \left[\ln \left(\frac{\Phi}{b} \right) \right]^2 \left(-4^{\frac{1}{1-k}} k^2 \Phi^{\frac{2k}{k-1}} + \Phi^2 k(k-1) \right) \right. \\ \left. - 4k(k-1)(d-1) \ln \left(\frac{\Phi}{b} \right) \left(4^{\frac{1}{1-k}} \Phi^{\frac{2k}{k-1}} - \Phi^2 \right) \right. \\ \left. + (d-2)(d-1)(k-1) \left(-4^{\frac{1}{1-k}} (k-1) \Phi^{\frac{2k}{k-1}} + \Phi^2 k \right) \right\}, \quad (2.12)$$

and the equations of motion (2.5) admit a solution where the scalar field and the metric are given as follows

$$\Phi(r) = b e^{ar}, \quad F(r) = r^2 \left(1 - \left[\frac{1}{4} b^2 e^{2ar} \right]^{\frac{1}{k-1}} \right).$$

3 Thermodynamics

The partition function for a thermodynamical ensemble is identified with the Euclidean path integral in the saddle point approximation around the Euclidean continuation of the classical solution [23]. The Euclidean and Lorentzian action are related by $I_E = -iI$ where the periodic Euclidean time is $\tau = it$. The Euclidean continuation of the class of metrics considered here is given by

$$ds^2 = N(r)^2 F(r) d\tau^2 + \frac{dr^2}{F(r)} + r^2 (dx_1^2 + dx_2^2 + \dots + dx_{d-2}^2).$$

In order to avoid conical singularity at the horizon in the Euclidean metric, the Euclidean time is made periodic with period β and the Hawking temperature T is given by $T = \beta^{-1}$. Here we are interested only in the static solution with a radial scalar field, and hence it is enough to consider a *reduced* action principle. This latter is given by

$$I_E = -\beta \Sigma_{d-2} \int_{r_h}^{\infty} \frac{N}{2k} (d-2) \frac{d}{dr} \left[r^{d-1} \left(1 - \frac{F}{r^2} \right)^k \right] \\ + \beta \Sigma_{d-2} \int_{r_h}^{\infty} N r^{d-2} \left\{ \left(\frac{1-4\xi}{2} \right) F(\Phi')^2 - \xi \Phi \Phi' \left(F' + \frac{2(d-2)}{r} F \right) - 2\xi \Phi \Phi'' F \right. \\ \left. - \frac{(d-2)\xi}{2r} F' \Phi^2 + \Phi^2 \left(-\frac{\xi}{2r^2} (d-2)(d-3) F \right) + U(\Phi) \right\} + B \quad (3.1)$$

where the radial coordinate r belongs to the range $[r_h, \infty[$ where r_h is the location of the horizon. Here, β stands for the inverse of the temperature and Σ_{d-2} corresponds to the compactified volume of the planar $(d-2)$ -dimensional base manifold. In the reduced action (3.1), B is a boundary term that is fixed by requiring that the Euclidean action has an extremum, that is $\delta I_E = 0$ which in turn implies that

$$\delta B = \beta \Sigma_{d-2} (\delta B_{\text{gravity}} + \delta B_{\text{matter}}) \quad (3.2)$$

where the first contribution $\delta B_{\text{gravity}}$ arises from the variation of the gravity part while the second one δB_{matter} comes from the matter source, and are given by

$$\begin{aligned} \delta B_{\text{gravity}} &= \int_{r_h}^{\infty} \frac{N}{2k} (d-2) \frac{d}{dr} \left[r^{d-1} \delta \left(1 - \frac{F}{r^2} \right)^k \right], \\ \delta B_{\text{matter}} &= \int_{r_h}^{\infty} 2\xi \Phi F r^{d-2} \delta \Phi' + \int_{r_h}^{\infty} \delta F \left[\frac{\xi(d-2)}{2} r^{d-3} \Phi^2 + \xi \Phi \Phi' r^{d-2} \right] \\ &\quad + \int_{r_h}^{\infty} \left\{ \delta \Phi \left[F \Phi' r^{d-2} (2\xi - 1) - \xi F' \Phi r^{d-2} \right] \right\}. \end{aligned} \tag{3.3}$$

In the grand canonical ensemble, the Euclidean action is related with the mass \mathcal{M} and the entropy \mathcal{S} by

$$I_E = \beta \mathcal{M} - \mathcal{S}. \tag{3.4}$$

For the solution (2.10), we have the following asymptotic variation behaviors

$$\begin{aligned} \delta F|_{\infty} &= -\alpha \xi^{\frac{1}{k-1}} a^{\alpha-1} (\delta a) \left[r^{\alpha+2} + \frac{b(\alpha-1)}{a} r^{\alpha+1} + \frac{b^2(\alpha-1)(\alpha-2)}{a^2} r^{\alpha} + \dots \right], \\ \delta \Phi|_{\infty} &= a^{\frac{(k-1)\alpha}{2}} (\delta a) \left[r^{\frac{(k-1)\alpha}{2}} + \frac{(k-1)\alpha b}{2a} r^{\frac{(k-1)\alpha-2}{2}} + \frac{(k-1)\alpha((k-1)\alpha-2)b^2}{4a^2} r^{\frac{(k-1)\alpha-4}{2}} + \dots \right], \\ \delta \Phi'|_{\infty} &= \frac{d}{dr} (\delta \Phi|_{\infty}), \end{aligned}$$

where for simplicity we have defined $\alpha = \frac{4\xi}{(k-1)(4\xi-1)}$. At the horizon r_h the variations read

$$\delta F|_{r_h} = -F'|_{r_h} \delta r_h, \quad \delta \Phi|_{r_h} = \delta \Phi(r_h) - \Phi'|_{r_h} \delta r_h.$$

It is intriguing that for the solution (2.10), the variation of the gravity part at the infinity exactly cancels at each order the variation of the matter at infinity, that is

$$\delta B_{\text{gravity}}(\infty) = -\delta B_{\text{matter}}(\infty),$$

and hence we have $B(\infty) = 0$. At the horizon, a simple computation yields

$$\delta B(r_h) = 2\pi \Sigma_{d-2} \delta \left[\left(1 - \xi \Phi^2(r_h) \right) r_h^{d-2} \right]$$

but as stressed before (2.11) the quantity between brackets vanishes. Hence, the boundary term vanishes identically, $B = 0$, and the identification of the mass and the entropy through (3.4) yields

$$\mathcal{M} = 0, \quad \mathcal{S} = 0.$$

It is then clear that the solutions obtained here have the particularity that their entropy is proportional to the lapse function evaluated at the horizon (2.11), and this quantity, by definition of the horizon, precisely vanishes. Consequently, the integration constant

a appearing in the solution can be naively interpreted as a sort of hair since it has not conserved charged associated to it. The temperature of the solution is given by

$$T = -\frac{\xi^{\frac{8\xi-1}{4\xi}}}{\pi a(k-1)(4\xi-1)} \left(\xi^{\frac{1-4\xi}{4\xi}} - b \right)^2, \quad (3.5)$$

and turns out to be positive provided that the nonminimal coupling parameter $\xi < \frac{1}{4}$ which is precisely the range where the solutions behaves asymptotically AdS. Interestingly enough, the temperature being inversely proportional to the hair a , this implies that a high temperature the hair will disappear as it is excepted in order to reproduce the phase diagram of the unconventional superconductors. It is also strange that the mass as well as the entropy of these solutions vanish identically.

4 Conclusions

Here, we have considered some particular Lovelock gravity theories indexed by an integer k whose coefficients are fixed by requiring the existence of a unique AdS vacuum with a matter source given by a self-interacting nonminimally coupled scalar field. We have shown that for each inequivalent Lovelock gravity theories there exists an appropriate choice for the self-interacting potential that permits to obtain black hole solutions for any arbitrary values of the nonminimal coupling parameter. The form of the potential involves six different terms where the first three terms exactly correspond to the potential that usually arises in different situations involving nonminimally coupled scalar fields. It will be interesting to explore in which context the k -depending part of the potential U_k (which concerns the last three terms in (2.7)) may emerge. The thermodynamics analysis of the solutions shows that the mass and the entropy of the solutions both vanish, and hence the integration constant a of the solutions can be viewed as a sort of hair. Such kind of black hole solutions with vanishing entropy and mass have also be found in the case of pure Lovelock gravity in [24]. In this reference, the authors consider a direct product $\mathcal{M}^m \times \mathcal{H}^n$ where \mathcal{H}^n is a negative constant curvature space, and the field equations of the m -dimensional space are trivially satisfied. Black hole solutions with zero mass and entropy have also be found in pure Lovelock gravity in even dimension [25] as well as in the case of Lifshitz black hole solutions of R^2 gravity [26]. We believe that, in all these cases, the most general solution has not been found and only particular solutions have been derived where some parameters have been switched out, and these latter may contain information about the mass of the solution. To illustrate this fact, we can take the example of the asymptotically AdS black solution found in [27] for new massive gravity in three dimensions. The metric solution is given by $F(r) = r^2 + br - 4GM$, and depends on two integration constants M and b where b is a sort of gravitational hair while M is related to the mass of the solution [28]. Hence if one would have only found the particular solution with $M = 0$, one would have concluded that the mass of the solution is vanishing.

We have also shown that the temperature T goes like $T \propto a^{-1}$ and hence at high temperature the hair should disappear as it is excepted in order to reproduce the phase diagram of the unconventional superconductors. It will be more than interesting to explore

more intensively the possible applications of the solutions derived here in the context of holographic superconductors. We end with the fact that for a coupling constant $b = 0$ and $\xi = \hat{\xi}_{k,d+1}$ as defined in (2.8), the potential reduces to a mass term and the solution becomes a stealth configuration [17, 18], that is a particular solution of the field equations (2.5) where both sides of the equations (gravity and matter) vanish identically

$$\mathcal{G}^{(k)} = 0 = T_{\mu\nu}.$$

In this particular case, the solution involves an additional integration constant which is due to the fact that the equation $T_{\mu\nu} = 0$ becomes invariant under the rescaling of the scalar field $\Phi \rightarrow A\Phi$, where A is a constant. In order to be self-contained, we present the stealth solution as found in [17, 18] for $b = 0$ and $\xi = \hat{\xi}_{k,d+1}$

$$ds^2 = - \left(r^2 - (ar)^{\frac{1-d}{k}} \right) + \frac{dr^2}{\left(r^2 - (ar)^{\frac{1-d}{k}} \right)} + r^2 d\vec{x}_{d-2}^2,$$

$$\Phi = Ar^{\frac{(k-1)(1-d)}{2k}}.$$

Because of this additional constant A , the quantity $(1 - \xi\Phi^2)$ evaluated at the horizon may not necessarily be zero, and consequently, the entropy may not be zero. We would like to further explore this issue concerning the thermodynamics properties of the stealth solutions arising in the context of nonminimally coupled scalar field.

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