

Photometric redshifts and clustering of emission line galaxies selected jointly by DES and eBOSS

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ABSTRACT

We present the results of the first test plates of the extended Baryon Oscillation Spectroscopic Survey. This paper focuses on the emission line galaxies (ELG) population targetted from the Dark Energy Survey (DES) photometry. We analyse the success rate, efficiency, redshift distribution, and clustering properties of the targets. From the 9000 spectroscopic redshifts targetted, 4600 have been selected from the DES photometry. The total success rate for redshifts between 0.6 and 1.2 is 71% and 68% respectively for a bright and faint, on average more distant, samples including redshifts measured from a single strong emission line. We find a mean redshift of 0.8 and 0.87, with 15 and 13% of unknown redshifts respectively for the bright and faint samples. In the redshift range $0.6 < z < 1.2$, for the most secure spectroscopic redshifts, the mean redshift for the bright and faint sample is 0.85 and 0.9 respectively. Star contamination is lower than 2%. We measure a galaxy bias averaged on scales of 1 and 10 Mpc/ h of 1.72 ± 0.1 for the bright sample and of 1.78 ± 0.12 for the faint sample. The error on the galaxy bias have been obtained propagating the errors in the correlation function to the fitted parameters. This redshift evolution for the galaxy bias is in agreement with theoretical expectations for a galaxy population with $M_B - 5 \log h < -21.0$. We note that biasing is derived from the galaxy clustering relative to a model for the mass fluctuations. We investigate the quality of the DES photometric redshifts and find that the outlier fraction can be reduced using a comparison between template fitting and neural network, or using a random forest algorithm.

Key words: Cosmology: observations – Surveys – target selection – DES – redshift

1 INTRODUCTION

With the development of new technologies, instruments improving in their performance, we are planning wider and deeper sur-

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veys with a volume of an order of four magnitudes higher than 20 years ago. Our surveys fall into two categories: spectroscopic redshift surveys for galaxy clustering and photometric redshift surveys for gravitational lensing, clusters, and clustering. Photometric and spectroscopic surveys mutually need each other in the sense that spectroscopic surveys need photometry to estimate galaxy properties such as colors, stellar masses, approximate redshifts, but also to efficiently select and understand its targets such as biasing. Photometric surveys need spectroscopic data to quantify the accuracy of photometric redshift. Both aspects are subjects of this paper.

We first present the characteristics of the emission line galaxy samples observed with the extended Baryon Oscillation Spectroscopic Survey¹, (hereafter eBOSS). These data are part of the eBOSS ELG target selection definition effort, undergone in October 2014. We designed different target selections based on Sloan Digital Sky Survey (SDSS) (Ahn et al. 2014), the South Galactic Cap u-band Sky Survey (SCUSS) (Jia et al. 2014), and the Dark Energy Survey², hereafter DES.

eBOSS is a spectroscopic survey using the BOSS spectrograph (Smee et al. 2013) at the Apache Point Observatory. It will cover 7500deg² in a six-year period starting Fall 2014. eBOSS aims at measuring the baryon acoustic oscillation feature at redshift higher than 0.6 extending the first measurement from the SDSS at lower redshift (Percival et al. 2009). eBOSS will use a mixture of targets to have a measurement at $z \sim 1$ using high redshift emission line galaxies (ELG), quasars between redshift one and two (Leistedt & Peiris 2014), and Lyman-alpha absorption quasars at redshift higher than two (Font-Ribera et al. 2014). In this paper, we present the results of two high redshift ELGs target selections and test plates spectroscopic redshift distributions using the DES photometry. We also study the bias for DES-based target selections using the DES Science Verification data. Two companion papers study other photometric selections (Delubac et al. 2015) in preparation, and luminosity functions (Comparat 2015) in preparation. Comparat (2015) in preparation gives full details about this test observations detailing the pipeline of redshift reduction.

Several papers study the survey design of efficient target selection of ELGs such as Comparat et al. (2013); Adelberger et al. (2004) as well as observations such as DEEP2 (Newman et al. 2013), VIPERS (Garilli et al. 2014) and Wigglez (Parkinson et al. 2012). DESI (Schlegel et al. 2011) and eBOSS are a new generation of surveys using the latest technologies in the field of spectroscopy. Using these new instruments allows to cover large sky areas selecting higher redshift and fainter targets than preceding surveys, increasing the statistical confidence in the measurement of cosmological parameters.

On the photometric side, DES is an ongoing photometric ground-based galaxy survey which started in autumn 2013. DES uses the brand new 2.2 deg² DECam instrument (Flaugher et al. 2015) mounted on the 4m Victor M. Blanco Telescope located at the Cerro Tololo Inter-American Observatory (CTIO) in Chile. It will cover 5000deg² after completion in five optical broad bands observing the southern sky. DES will use cosmic shear, cluster counts, large scale structure measurements and supernovae to reach very competitive measurement of the Universe growth rate and dark energy.

One of the main limitations in ongoing and upcoming dark energy

photometric surveys are the access to the radial dimension, the redshift. Surveys such as DES, LSST, Euclid, have between 5 to 8 photometric broad bands. Broad band photometric redshifts (z_{ph}) have an accuracy limited by the filter resolution (Jouvel et al. 2011). In the optimal case of a space-based survey with 16 filters covering from UV to infrared, photometric redshifts reach a precision of 0.03 (Jouvel et al. 2014), in the case of crowded fields, with cluster interstellar light. Photometric redshifts precision is the unclipped rms of the redshift measurement error, the difference between the measured and true redshift. Depending on the survey configuration such as ground-based, space-based, number of broadband filters, pixel size and exposure time, the photometric redshifts accuracy will go from a bit better than 0.03 and degrade. DES reaches a precision of 0.08 (Sánchez et al. 2014) but this also depends on the color, magnitude, size, and redshift of the galaxy population considered. If photometric redshifts are used to measure galaxy clustering, then the resulting estimate of the dark matter power spectrum becomes biased at small scales because the density fluctuation traced by galaxies appear to be smoothed by the photometric redshift error. The bluest galaxies show a spectrum with higher degeneracies in color space allowing some high redshift galaxies to be misplaced at low redshift, and conversely. Dark Energy Equation of State constraints, which rely on photometric redshift information (like weak-lensing and cluster mass function estimates) can be severely affected by (unaccounted) outliers (Bernstein & Huterer 2009). In this paper we use the eBOSS test plates redshifts to investigate possible ways to reduce the outlier fraction.

In section 2.2, we present the definition, efficiency and spectroscopic characteristics of our DES photometrically selected sample, based on the DES Science Verification data (Melchior et al. 2014; Banerji et al. 2015; Sánchez et al. 2014). Using the eBOSS spectroscopic redshifts of the test observation we look at the photometric redshift accuracy of the DES survey and outlier removal in section 3. In section 2.3, we show the results of our systematics studies on the year one DES data which is an order of magnitude larger in area than the test plates presented in this paper. In section 4 we focus on the clustering properties of the spectroscopic samples and the impact of photometric redshift precision in the clustering from the point of view of DES. Assuming that the DES photometric redshift distribution are well calibrated, we study the impact of possible biases coming from the bins boundaries, due to photometric redshifts. Finally we present the conclusions on the characterization, suggestions for the improvement of the photometric DES target selection, and the outlier removal techniques.

2 DES AND EBOSS DATA

The eBOSS fields selected for the ELG target selection campaign were chosen to overlap with CFHTLS-W1. CFHTLS has four wide and four deep 1deg² fields in u^*, g', r', i', z' bands. The W1 field is the biggest of the wide fields with 19deg² and 80% completeness depth of $i' < 24.5$. Coupon et al. (2009) computed the CFHTLS photometric redshift accurate to 3-4% up to $i' < 22.5$ calibrated with VVDS (Le Fèvre et al. 2005), DEEP2 (Newman et al. 2013) and zCOSMOS (Lilly et al. 2007) spectroscopic surveys. These data are publicly available. This field has also been imaged by the SDSS survey³. The photometry from SDSS is a 7500 deg² with 95%

¹ <https://www.sdss3.org/future/eboss.php>

² <http://www.darkenergysurvey.org/>

³ <http://www.sdss.org/data>

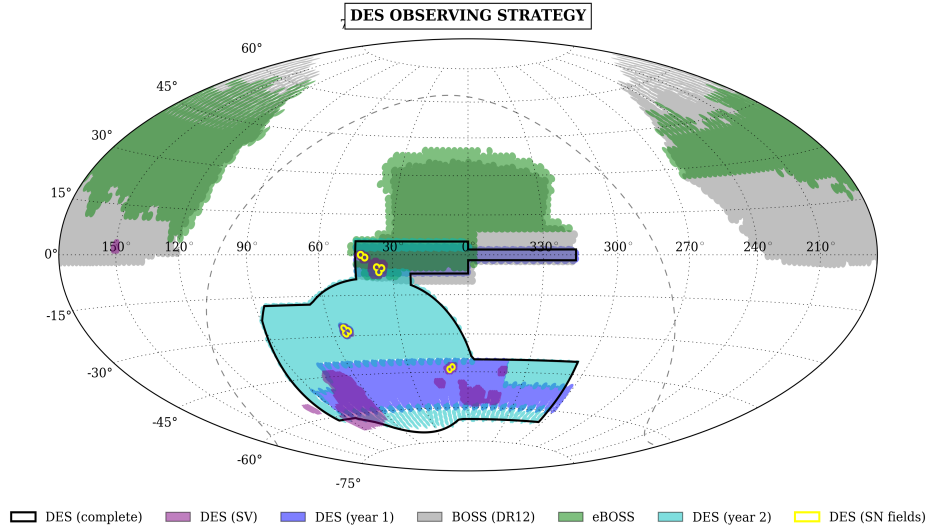


Figure 1. Footprints of DES, BOSS, and eBOSS. Coordinates are RA and Dec in deg.

completeness depth of $u, g, r, i, z = 22.0, 22.2, 22.2, 21.3, 20.5$ (Abazajian et al. 2009).

eBOSS and DES have an overlap of 500 deg^2 over Stripe82 which will yield a minimum number of 60000 spectra. Since DES photometry is deeper than SDSS, we optimise the target selection to reach fainter targets at higher redshift and lower contamination from low redshift galaxies. Abdalla et al. (2008) showed that a neural network can pick up strong emission lines from the broad-band photometry of a galaxy using DEEP2 and SDSS data. Further optimisation has been done applying this method to a DESI-like survey in Jouvel et al. (2014). Jouvel et al. (2014) shows that one reaches a higher success rate using neural networks target selections than boxes in color space. A related work used a Fisher discriminant method to investigate improvements in the eBOSS target selection (Raichoor 2015). We note that success rate here is the percentage of objects for which we have a very secure galaxy redshift. The completeness of the target selection in terms of galaxy population would be the subject of another paper.

2.1 DES Science Verification data

The first phase of the DES survey consisted of various tests and improvements in the data acquisition, instrument calibration and data processing, which resulted in a first well defined source catalog, the Science verification data, hereafter SVA1. Scientific results from Bonnett et al. (2015); Rozo et al. (2015); Crocce et al. (2015); Banerji et al. (2015); Sánchez et al. (2014); Melchior et al. (2014) and others show the very good quality of the SVA1 data, going at the nominal depth of the DES survey. In Figure 1 we show the footprint of the DES, BOSS and eBOSS surveys along with the DES year one data, and the SVA1 data.

To define the target selection of the eBOSS ELG sample, explained in the next subsection, we used an area of 9.2 deg^2 from the SVA1 data. Figure 2 shows the depth over the 9.2 deg^2 SVA1 footprint in g, r, z used for the eBOSS observations. Limiting magnitudes are defined by the flux in a 4 arcseconds aperture above 10σ , computed from the DES images using the Mangle software (Swanson et al. 2012). Details about the resulting spectroscopic sample will be presented in subsection 2.3.

In section 3 we explore the photometric redshifts (z_{ph}) for our

sample. We use the ANNz2 code (Sadeh 2015), which is a new version of ANNz (Collister & Lahav 2004). We also show photometric redshifts from the template fitting code Le Phare (Ilbert et al. 2006, 2009). We computed zeropoints corrections with the ELG samples and applied them to the Le Phare results. Except from this, the Le Phare parameters used here are the same than the one used in Sánchez et al. (2014). The random forest code TPZ⁴ (Carrasco Kind & Brunner 2013) is used to identify catastrophic redshifts. Photo- z results from these codes have been previously studied for DES in Sánchez et al. (2014). They use DES-SVA1 data trained with a spectroscopic redshift (z_{sp}) sample independent of the eBOSS data. Sánchez et al. (2014) training sample contains 9000 z_{sp} from various spectroscopic surveys. ANNz2 includes a weighting of the galaxies during the training procedure which is a function of magnitude, color and spectroscopic redshifts to correct for the sample difference between the photometric and spectroscopic data.

To separate stars and galaxies we use a combination of SExtractor parameters combined to have an optimum purity/completeness, hereafter `modest_class`. `Modest_class` stars are defined in Table 1

Note that we add selection criterion at $g < 23$ for the `modest_class` classification.

2.2 eBOSS ELG spectroscopic targets

We used three tiles from the SVA1 data on CFHTLS-W1 which we observed in 8 eBOSS plates. With a one hour exposure we reached z a total of 5705 spectra. We investigated three different target selection schemes, see Table 2. SDSS-SCUSS targets are dispatched over 51 deg^2 . DES bright and faint targets are dispatched over 9.2 deg^2 . SDSS-SCUSS is a combination of an SDSS only and a SDSS-SCUSS target selection as can be seen in Table 2. The DES faint selection has been optimized to reach redshifts between 0.7 and 1.5. This latter selection has been designed for the DESI survey and reaches higher redshifts galaxies than eBOSS is

⁴ <http://lcdm.astro.illinois.edu/code/mlz.html>

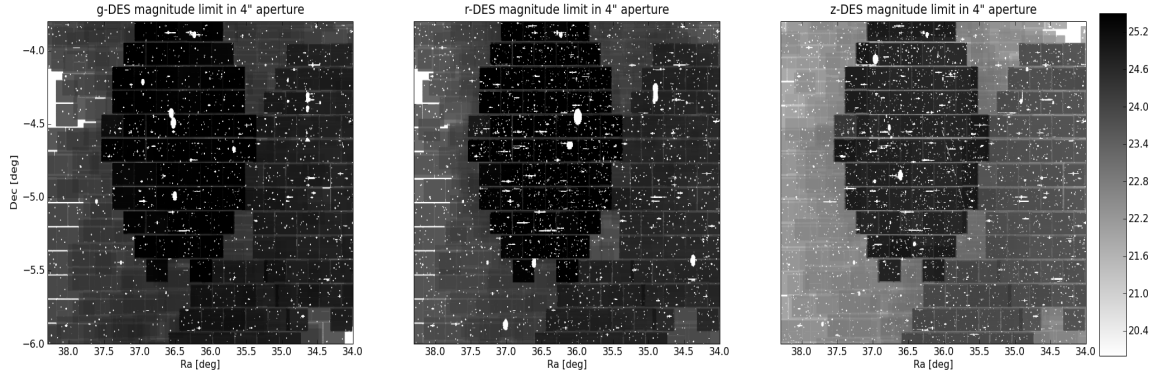


Figure 2. Depth of the g,r,z-band of the DES Science Verification data in the 9.2 deg^2 of the eBOSS test plates. The depth has been computed with the Mangle software and corresponds to a 10σ magnitude in an aperture of 2 arcsec. Find more details in section 4.

Table 1. Modest_class stars. Note that we add selection criterion at $g < 23$ for the modest_class classification. Magnitudes are the SExtractor (Bertin & Arnouts 1996).

$$\left\{ \begin{array}{l} \text{FLAGS_I} \geq 3 \\ \text{OR} \\ \text{FLAGS_I} \geq 3 \quad (\text{SPREAD_MODEL_I} + 3 * \text{SPREADERR_MODEL_I}) < 0.003 \quad (\text{SPREAD_MODEL_I} + 3 * \text{SPREADERR_MODEL_I}) > -0.003 \end{array} \right. \quad \left. \begin{array}{l} \text{CLASS_STAR_I} > 0.3 \\ \\ \\ \text{MAG_AUTO_I} < 18.0 \quad \text{MAG_PSF_I} < 30.0 \end{array} \right\}$$

Table 2. The three eBOSS ELG selections. *eg, er, ei, ez* are photometric uncertainties of the g,r,i,z bands. Magnitudes are the SExtractor (Bertin & Arnouts 1996) detmodel DES magnitudes for the bright and faint selections and SDSS/SCUSS model magnitudes for the SDSS-SCUSS selections.

$$\left\{ \begin{array}{l} \text{DES bright:} \quad 20.5 < g < 22.8 \quad -0.7 < g-r < 0.9 \quad 0 < r-z < 2 \quad r-z > 0.4 * (g-r) + 0.4 \\ \text{DES faint:} \quad g > 20.45 \quad r < 22.8 \quad 0.28 < r-z < 1.58 \quad g-r < 1.15 * (r-z) - 0.2 \quad g-r < 1.45 - 1.15 * (r-z) \\ \text{SDSS-SCUSS:} \quad eg < 0.6 \quad 20 < g < 23 \quad r < 22.5 \quad i < 21.6 \quad 21 < U < 22.5 \quad r-i > 0.7 \quad i-u > -3.5 * (r-i) + 0.7 \\ \quad \quad \quad er < 1 \quad 21 < g < 22.5 \quad r < 22.5 \quad i < 21.6 \quad g-r < 0.8 \quad r-i > 0.8 \\ \quad \quad \quad ei < 0.4 \end{array} \right. \}$$

aiming at (Schlegel et al. 2011). eBOSS is aiming at galaxies between a redshift of 0.6 and 1.2. Higher redshifts will be explored using AGNs. This paper studies the DES ELG target selections for eBOSS. The two other selections using SDSS and SCUSS data will be presented in our companion paper Comparat (2015); Delubac et al. (2015); Raichoor (2015). On the DES selections we apply the star-galaxy separation explained in section . using a combination of SExtractor output optimized for the DES data. We note that we do not expect much contamination by stars when designing ELG target selection as shown in Adelberger et al. (2004). We remove the fake detections by applying selection criteria in g,r,z DES bands of $\text{MAG_APERTURE} - \text{MAG_DETMODEL} < 2$. Figure 3 shows the photometric redshifts distribution in cyan solid line of the galaxies in the DES 9.2 deg^2 field used to optimise the eBOSS target selections. The cyan solid line shows all galaxies at $g < 23$ with a photometric redshift between 0.5 and 1.5. We show two of the DES based target selections in dash-dotted blue and dashed green lines. The red dotted line shows an SDSS-based selection which we name SDSS-SCUSS, detailed in the next section. The magenta triangles show the outliers. We note that there is color space where there seem to be a higher percentage of outliers, especially for the DESI selection. In section 3.3, we present a first attempt to find color-magnitude boxes where we find a higher percentage of outliers.

The eBOSS ELG observations are presented in Table 3. We show the number of targets selected and observed in the eBOSS

test plates. SDSS-SCUSS do not show the total number of targets observed but the one for which we find a match with DES photometry. For representative statistics about the SDSS and SCUSS selections, please refer to Comparat (2015); Delubac et al. (2015); Raichoor (2015). We show the percentage of z_{sp} with a secure redshift “secure” for which we find at least two lines with a low signal-to-noise detection, or one line and a 10σ continuum detection for the redshift measurement. The “1line” were measured from a single line with at least 3σ detection without continuum information. They have a higher failure rate since a line confusion can happen between $[H\alpha]$ and $[OII]$. The “unknown” are spectra for which we couldn’t find a redshift. ‘ $0.6 < z < 1.2$ ’ shows the percentage of targets with secure redshifts in the desired redshift range: $0.6 < z < 1.2$. ‘ $0.6 < z < 1.2*$ ’ includes the ‘1line’ z_{sp} in the percentage of targets with secure redshifts. \bar{z} and $< [OII] >$ are respectively the mean eBOSS z_{sp} and $[OII]$ flux using secure redshifts. Nstars is the number of stars.

DES gives the highest success rate with 72% of very secure redshift, 12% of one line detected redshift, and 15% of non-identified redshift. DES faint selection has a slightly lower success rate of 68% of very secure redshifts including 20% of one line z_{sp} . Table 3 shows the success rate as a function of DES g-band magnitude. The DES faint selection has been designed to target fainter and higher redshift galaxies which explains the slightly lower success rate when compared to DES bright selection. In section 2.3,

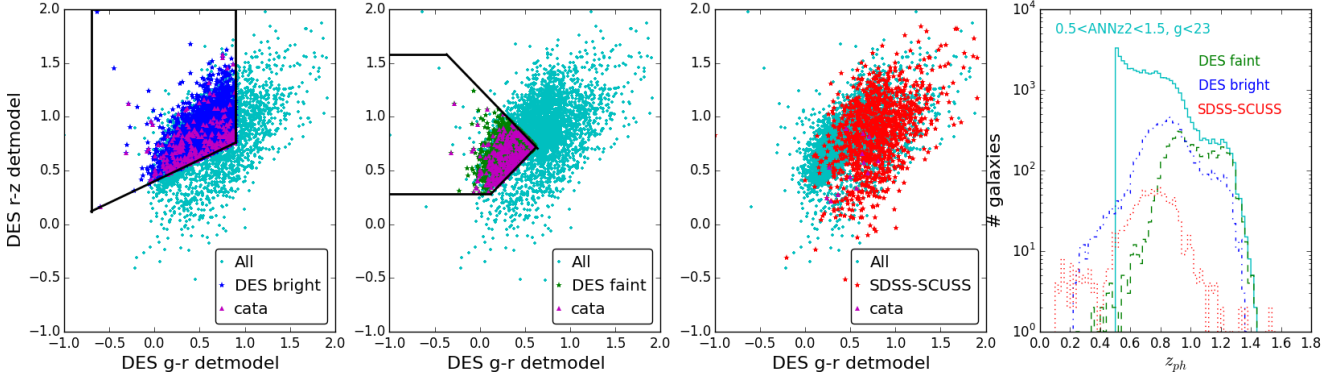


Figure 3. The three panels from the left-hand side show the DES detmodel grz colors of the eBOSS plates, showing the ELG selection functions. The magenta triangles labeled ‘cata’ shows catastrophic redshifts as defined in section 3. The furthest right panel shows the photometric redshift distribution of the eBOSS targets selected with the DES-SVA1. Photo-z are computed in using ANNz2 on the DES SVA1 data for the DES selections. For SDSS-SCUSS selection, we used CFHTLS photo-z which we matched with SDSS photometry. The median uncertainties on colors $g - r$ and $r - z$ for the different selections is less than 4%.

we apply the DES bright selection to the year one DES data and show the results of our systematics studies.

Figure 4 shows a comparison between our spectroscopic redshift measurement with other spectroscopic redshift surveys available on the CFHT-LS field: VVDS_DEEP (Le Fèvre et al. 2005), DEEP2 (Newman et al. 2013), GAMA⁵, VIPERS (Garilli et al. 2014), OzDES (Yuan et al. 2015). We used only the most secure redshift for both eBOSS and other z_{sp} surveys. We matched catalogues at less than one arcsec.

With the VIPERS survey, using the quality flags 3.1, 3.2, 3.5, 4.1, 4.2 and 4.5, we find 4 galaxies out of 243 with a discrepant redshift. For OzDES, we found 4 galaxies with discrepant redshift out of 95. However, restricting to the flag 4 only, we found no discrepant redshift between OzDES and eBOSS. Similarly, we found no discrepant redshift with the 23 galaxies in VVDS. If we restrict the match between catalogues at less than 0.1 arcsec, we don’t find any disagreement between VIPERS and eBOSS for the flags 3 and 4 redshifts. Comparat (2015) in preparation gives a full detailed study of the eBOSS redshift measurement pipeline and tests and finds an agreement with VIPERS at less than 1%. We conclude that eBOSS redshift measurement are very reliable.

2.3 eBOSS DES ELG selection characteristics

As mentioned in the introduction, DES and eBOSS footprint overlap on about 500deg^2 in the Stripe82 region. The early release of the DES year one data, hereafter Y1A1, is on the Stripe82 field with 152deg^2 . We use the Y1A1 data to look at possible systematics from the photometric selections. In order to have reliable measurement of the galaxy power spectrum in eBOSS we need a density variation of maximum 15% over the whole survey area as discussed in Ross et al. (2012) and Dawson (2015) in preparation. We use Healpix⁶ to produce maps of the eBOSS galaxy target selection and systematics which have the biggest impact on the power spectrum measurement such as stellar density, Galactic extinction, survey depth, airmass (Ross et al. 2012). We used the DES bright selection for the galaxy density maps using a pixelisation of

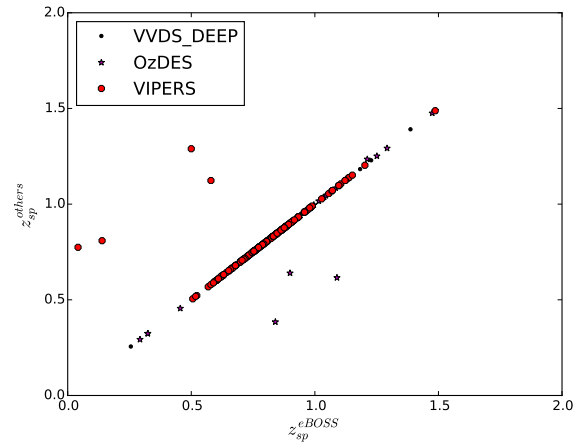


Figure 4. Spectroscopic redshift from eBOSS compared to other z_{sp} surveys. We use only the secure z_{sp} .

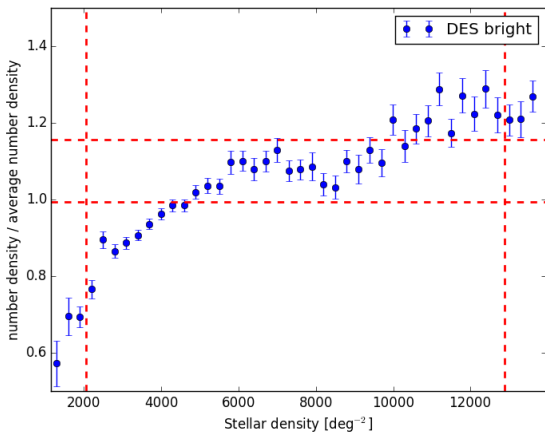
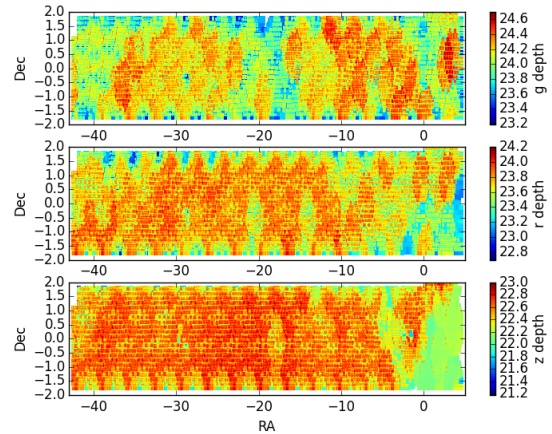
6.87arcmin^2 (NSIDE=512). Note that we add a selection criterion: $g < 23$ for the modest_class classification. We are aiming to test the impact of the stars density variation with the eBOSS bright galaxies density variation. The density of stars at lower magnitude will not impact our test. The number density of stars vary between 76 and 15000 with a mean of 5890 stars/ deg^2 . The mean number density of galaxies is 737 gal/ deg^2 with variations between 19 to 2690 gal/ deg^2 . In Figure 5, we show the galaxy density fluctuation as a function of the star density using modest_class star classification. The galaxy density variation caused by stars and their haloes is less than 15% across the the Y1A1 Stripe 82 survey area. We conclude that the photometric contamination of star haloes does not have a significant impact on our target selection. Similarly, in Figures A1, we looked at the variation of target density as a function of depth, Galactic extinction, airmass across the Y1A1 Stripe82 and conclude that our target selection will be within requirement across the Stripe82 DES footprint.

⁵ <http://www.gama-survey.org>

⁶ <http://healpy.readthedocs.org/en/latest/>

Table 3. Number of targets for the four eBOSS selections. We use the DES ‘g’ band to do magnitude selections.

	DES bright	DES faint	SDSS-SCUSS	DES bright \cap faint	
$20.5 < g < 22$	Selected	953	445	-	220
	Dens. Selected /deg ²	69	32	-	24
	Observed	557	254	206	199
	secure(%)	88.0	85.0	77.2	87.4
	l1line(%)	1.3	3.5	1.0	2.5
	unknown(%)	10.8	11.4	21.8	10.1
	$0.6 < z < 1.2$ (%)	60.9	66.5	58.3	71.4
	$0.6 < z < 1.2^*$ (%)	61.4	67.7	58.3	72.4
	\bar{z}	0.68	0.8	0.65	0.8
	$\langle [OII] \rangle$	1.8	2.54	3.91	2.65
	Nstars	21	13	25	5
$22 < g < 23$	Selected	6762	7838	-	2158
	Dens. Selected /deg ²	491	570	-	239
	Observed	3103	2158	1049	1274
	secure(%)	70.6	64.0	74.5	67.1
	l1line(%)	13.4	22.3	4.6	21.3
	unknown(%)	16.0	13.6	21.0	11.6
	$0.6 < z < 1.2$ (%)	64.8	55.7	62.7	60.4
	$0.6 < z < 1.2^*$ (%)	73.5	68.1	66.0	72.0
	\bar{z}	0.83	0.88	0.71	0.9
	$\langle [OII] \rangle$	1.24	1.47	0.97	1.71
	Nstars	51	35	24	16
$20.5 < g < 23$	Selected	7716	8283	-	2378
	Dens. Selected /deg ²	561	602	-	264
	Observed	3660	2412	1255	1473
	secure(%)	73.3	66.3	74.9	69.9
	l1line(%)	11.6	20.4	4	18.7
	unknown(%)	15.2	13.4	21.1	11.4
	$0.6 < z < 1.2$ (%)	64.2	56.8	62	61.8
	$0.6 < z < 1.2^*$ (%)	71.6	68.0	64.7	72.0
	\bar{z}	0.8	0.87	0.7	0.88
	$\langle [OII] \rangle$	1.34	1.61	1.47	1.87
	Nstars	72	48	49	21

**Figure 5.** Density fluctuation of galaxies as a function of the stellar density. The two horizontal and vertical red axes show respectively the 5 and 95% of the star density distribution and 15% around the error weighted mean galaxy density fluctuation for the eBOSS target selection. The error weighted mean is computed using 94% of the galaxy population, pruning from the extremes low and high 2% of the density regions. Errorbars show the standard deviation of the galaxy number density over the healpix pixels, that we used as weight in the mean calculation.**Figure 6.** Depth of the DES year one data on Stripe 82 region in the g,r,z bands.

3 PHOTO-Z FROM THE EBOSS OBSERVATIONS

3.1 Photo-z of the four eBOSS selections

We matched positions of the ANNz2 photometric redshifts catalogue (Sánchez et al. 2014) with positions of eBOSS targets with spectroscopic redshifts. We show photometric redshifts of the four

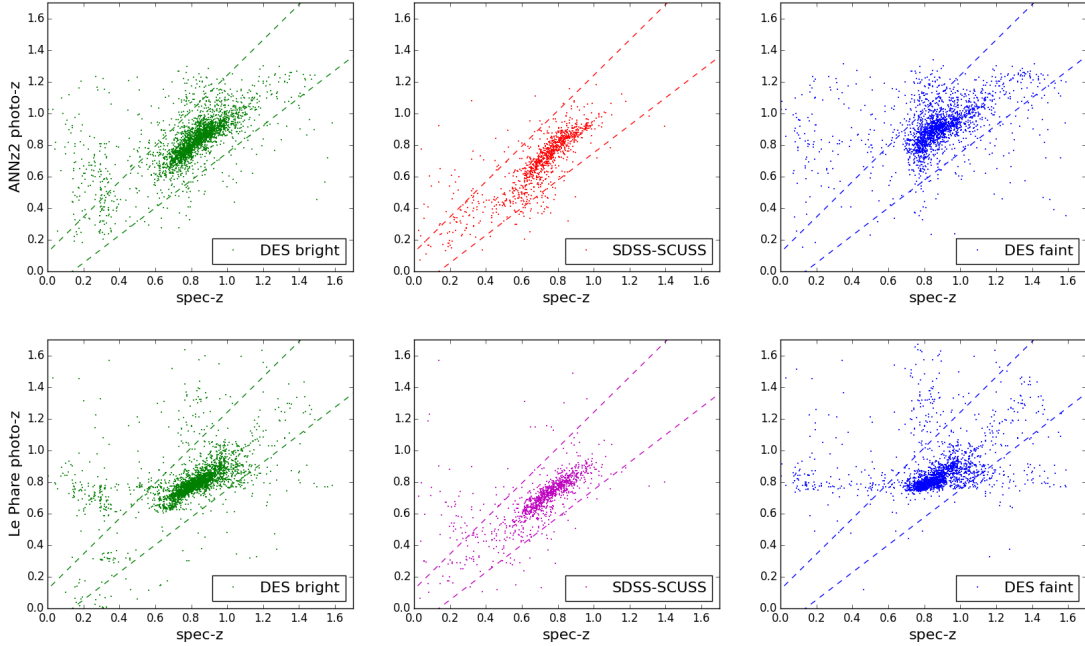


Figure 7. z_{ph} vs z_{sp} of the four eBOSS target selections. We use ANNz2 code for the top row and Le Phare for the bottom row, on the five band DES-SVA1 photometry. The dashed line for the top and bottom panel shows the expected DES accuracy of $|z_{ph} - z_{sp}| = \sigma_{DES} * (1 + z_{sp})$ where $\sigma_{DES} = 0.12$.

eBOSS selections in Figure 7, using a comparison with the secure redshifts only. Note that Sánchez et al. (2014) dealt with the full DES galaxy population, while the eBOSS emission-line galaxy sample considered here is a subset of the galaxy population for which photo-z’s are prone to larger errors. The SDSS-SCUSS selection has better photometric redshifts than the DES bright and faint selections as shown in Table 4. SDSS-SCUSS are targeted using the SDSS photometry which is shallower than the DES photometry. SDSS-SCUSS targets are redder in g-r than DES targets as shown in Figure 3. Galaxies with a large difference between g and r have a stronger Balmer or 4000Å break making the redshift measurement easier for $z < 1$ galaxies. Indeed the SDSS-SCUSS selection has criteria in r-i and u-i/g-r excluding the selection of galaxies with SED-power-law, which are difficult to locate in redshift space. DES selections will then have a higher catastrophic rate as shown in Figures 3, 7 and Table 4. Table 4 shows the mean and median z_{ph} , standard deviation, NMAD and outliers fraction of the $z_{ph} - z_{sp}$ distribution as a function of g-band magnitude for the four selections. NMAD is the normalised median absolute deviation defined as $1.48 * \text{median}(|z_{ph} - z_{sp}|) / (1 + z_{sp})$. NMAD is a calculation of the $z_{ph} - z_{sp}$ dispersion reducing the deviations induced by outliers as explained in Ilbert et al. (2009).

For comparison we also obtained photometric redshifts with the Le Phare code (Arnouts et al. 1999; Ilbert et al. 2006). We used the template library developed for the COSMOS observations in Ilbert et al. (2009). It has 31 templates from elliptical to starbursts galaxies. We apply extinction to the bluest templates using the Calzetti law (Calzetti et al. 2000) with extinction values of $E(B - V) = [0.1, 0.2, 0.3]$. We enable the Le Phare in-built redshift prior calibrated with the VVDS observations (Le Fèvre et al. 2005). DES observations have the same depth as VVDS which justify the use of this prior. Figure 7 shows Le Phare and ANNz2 photometric redshifts as a function of the eBOSS spectroscopic redshifts. ANNz2 and Le Phare photometric redshifts have simi-

lar performances, although Le Phare has a tendency to aggregate galaxies at z_{ph} 0.8. This is a feature caused by the discretisation of the redshift-template space encoded in most template fitting photometric redshifts software.

3.2 Removing catastrophics: template fitting vs neural network

In the context of the DES survey, we investigate possible ways to remove catastrophic redshifts. ELG are the most difficult galaxy population to find an accurate z_{ph} . ELG are the source of the high percentage of outliers. We take advantage of the eBOSS sample to look at possible ways of calibrating and removing a part of the outliers fraction. A first solution has been proposed in Newman et al. (2013). Comparing template fitting methods and machine learning helps at pruning outliers. Figure 8 top panel shows an example of this method with the DES-eBOSS data. Top panel of Figure 8 shows the density of $LePhare - z_{sp}$ versus $LePhare - ANNz2$. Accurate photometric redshifts are located inside the black lines. The red lines show a possible template fitting vs machine learning criterion at $|ANNz2 - LePhare| = 2 * \sigma_{DES}$ where $\sigma_{DES} = 0.12$ is the expected accuracy for the DES survey. Selecting the galaxies inside the red lines is a possible way to prune from outliers. It will however remove some galaxies for which we have a good photometric redshift. There is a tradeoff between removing more outliers in moving the red lines and not losing too many galaxies for the clustering analysis. The bottom panel of Figure 8 shows the purity and completeness trend while changing the $2 * \sigma_{DES}$ value to the value of the x-axis. The purity shows the percentage of catastrophic redshifts that the selection criteria will remove. Catastrophic redshifts are defined as $|LePhare - z_{sp}| > 2\sigma_{DES}$. The completeness shows the percentage of galaxies left in the sample. A selection criterion at $2\sigma_{DES}$ removes about 30% of outliers and leaves 92% of the galaxy sample. Following this idea, a more de-

Table 4. ANNz2 photometric redshifts results for the four eBOSS selections. The DESz selection corresponds to the DES targets selected at $0.6 < z_{sp} < 1.2$.

		DES bright	DES faint	SDSS-SCUSS	DESz bright	DESz faint	SDSS-SCUSSz
$20.5 < g < 22$	z_{ph}^-	0.72	0.8	0.63	0.8	0.83	0.71
	median(z_{ph})	0.74	0.82	0.67	0.79	0.85	0.72
	$\sigma[z_{ph} - z_{sp}]$	0.21	0.3	0.12	0.14	0.2	0.09
	NMAD[$z_{ph} - z_{sp}$]	0.07	0.11	0.04	0.05	0.08	0.03
	outliers	101	71	10	43	40	5
$22 < g < 23$	z_{ph}^-	0.85	0.93	0.7	0.85	0.92	0.75
	median(z_{ph})	0.85	0.92	0.72	0.85	0.92	0.75
	$\sigma[z_{ph} - z_{sp}]$	0.15	0.22	0.11	0.1	0.13	0.09
	NMAD[$z_{ph} - z_{sp}$]	0.04	0.05	0.04	0.03	0.04	0.03
	outliers	187	223	37	83	104	21
$20.5 < g < 23$	z_{ph}^-	0.83	0.91	0.69	0.84	0.91	0.74
	median(z_{ph})	0.84	0.92	0.72	0.84	0.91	0.75
	$\sigma[z_{ph} - z_{sp}]$	0.17	0.23	0.11	0.11	0.14	0.09
	NMAD[$z_{ph} - z_{sp}$]	0.04	0.06	0.04	0.03	0.05	0.03
	outliers	288	294	47	126	144	26

tailed work has also been pursued in Carrasco Kind & Brunner (2014).

3.3 Removing catastrophics using Random Forest

We investigated a new method to remove outliers using the TPZ code which is part of the MLZ framework (Carrasco Kind & Brunner 2013). TPZ is a random forest algorithm which, in this case, divides the color space in branches, forming multi-dimensional color boxes. We used a random half of the eBOSS spec-z as a training set and the other half as testing set. We give TPZ the DES magnitudes and colors. We looked at the percentage of catastrophics for each of the branches. The top panel of Figure 9 shows the number of galaxies as a function of the number of catastrophic redshifts for each of the branches. The density corresponds at the number of branches. We observe that some branches have a high percentage of catastrophic redshift for a high number of galaxies in the box. Same as in section 3.2, there is a tradeoff between the percentage of outliers one can remove and the percentage of galaxies which are left in the sample. The bottom panel of Figure 9 shows the purity and completeness while using this method. Lines show the number of galaxies left in the sample as a function of the percentage of outliers selection criterion value used to trim the branches. For example, if we exclude the boxes which have more than 40% of outliers, we are left with about 85% of galaxies and 50% of outliers, depending on which photometric redshift catalogue is used to define the outliers. In black solid, blue dashed, and green dashed-dotted, respectively we show the results for TPZ, LePhare, and ANNz2. The axis on the right handside shows the percentage of galaxies left. We note that this method relies on having a spectroscopic sample (i.e., the eBOSS test plates), and that this sample then defines the good color boxes which may then be used in the future (if planned) to select eBOSS samples with fewer catastrophic outliers. This method could also be applied to decrease the outlier fraction in the DES photo-z main sample, assuming a representative training set.

4 CLUSTERING PROPERTIES

In this section we measure the clustering properties of the DES galaxy target selections. We analyze the faint and bright selections separately and estimate the galaxy bias of these samples. Further clustering analyses for different eBOSS target selections will be

	bright		
	Number (purity)	mean density	mean redshift
z_{sp}	2613 (100%)	284.02 gal/sqdeg	0.855 (0.855)
ANNz2	2902 (86.66%)	315.43 gal/sqdeg	0.866 (0.902)
LePhare	3038 (84.10%)	330.22 gal/sqdeg	0.811 (0.894)
	faint		
	Number (purity)	mean density	mean redshift
z_{sp}	2139 (100%)	232.50 gal/sqdeg	0.901 (0.901)
ANNz2	2582 (79.43%)	280.65 gal/sqdeg	0.928 (0.994)
LePhare	2662 (77.23%)	289.35 gal/sqdeg	0.841 (1.004)

Table 5. Properties for the bright and faint sample including the different photometric redshifts catalogues properties. The neural network redshift code seems to have a better performance. ANNz2 code produces counts and number densities that more closely resemble those from the sample with spectroscopic redshifts.

presented in a separate work in preparation. Throughout this analysis we assume a flat Λ CDM + ν (one massive neutrino) cosmological model based on Planck 2013 + WMAP polarisation + ACT/SPT + BAO, with a total matter density relative to critical $\Omega_m = 0.307$ (Planck Collaboration et al. 2014)).

Tables 5 show the number of galaxies, mean density and redshift for the bright and faint samples considered in the clustering analysis. The spectroscopic catalogue has been selected using the true spectroscopic redshift in a given redshift range, while the ANNz2 and LePhare catalogues are selected from their respective photometric redshift estimations, for galaxies in the range $0.6 < z_{ph} < 1.2$, but using their true redshift in order to compute distances and estimate the monopole of the two-point spatial correlation function (see below). The purity of each sample was calculated as the ratio of the number of true galaxies with spectroscopic redshifts in the given range and the number of galaxies in the spectroscopic sample. The mean redshift is calculated with the photometric redshifts, and in parenthesis, with the spectroscopic redshift of the same sample. By comparing the mean z_{ph} and z_{sp} , we get an estimate of the bias for each sample. The effective area of the footprint is 9.2 deg^2 . These numbers were obtained after the catalog was pruned by the angular mask, as detailed below.

The redshift distribution for the spectroscopic sample is shown in Figure 7.

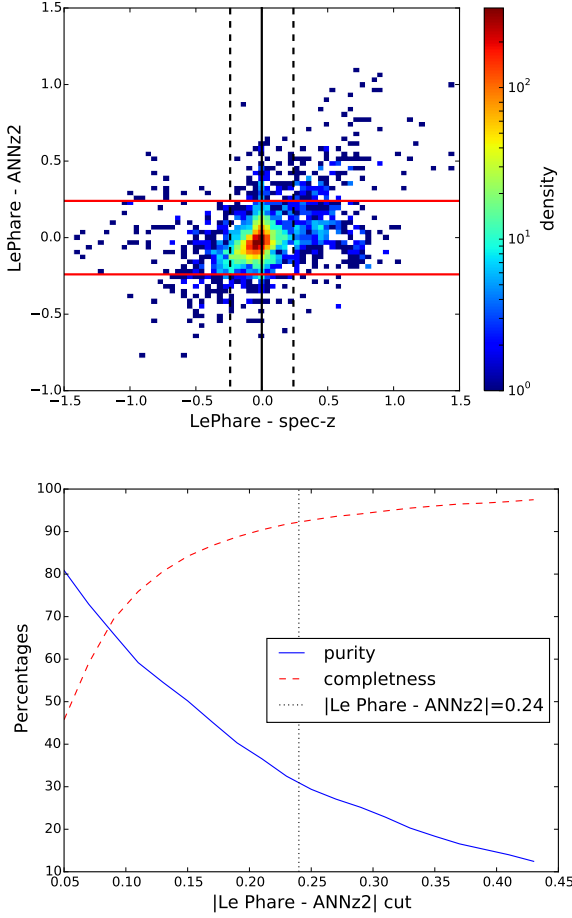


Figure 8. Top panel shows the z_{ph} improvement while doing a selection criterion on the difference between template fitting and machine learning z_{ph} . The bottom panel shows the purity-completeness while changing the selection criterion. The bottom panel shows how well this method would perform.

4.1 Random fields

We use the Mangle mask of the DES survey (Figure 2) in the observed field to create a random catalog, sampling the footprint with the same depth and angular distribution than the ELG selections. We use as reference observations the g band, where the magnitude limit is imposed to create the target selection (see target selection in 2.2). The magnitude limit distribution for the area analysed is shown in the first panel of Figure 2 in units of magnitude in a 2 arcseconds aperture at 10σ .

We compute the random sample considering the different magnitude limits of the area, applying a limiting aperture magnitude selection in the mask at $23.4 < mag < 25.8$. The final area amounts to 9.2 deg^2 .

As we observe different galaxy densities depending on the survey depths, we take this information into account when generating the random catalogs. In order to generate non-uniform random catalogs according to DES depths we have applied the following methodology:

- (i) Create a uniform random catalog according to the galaxy catalog angular footprint.

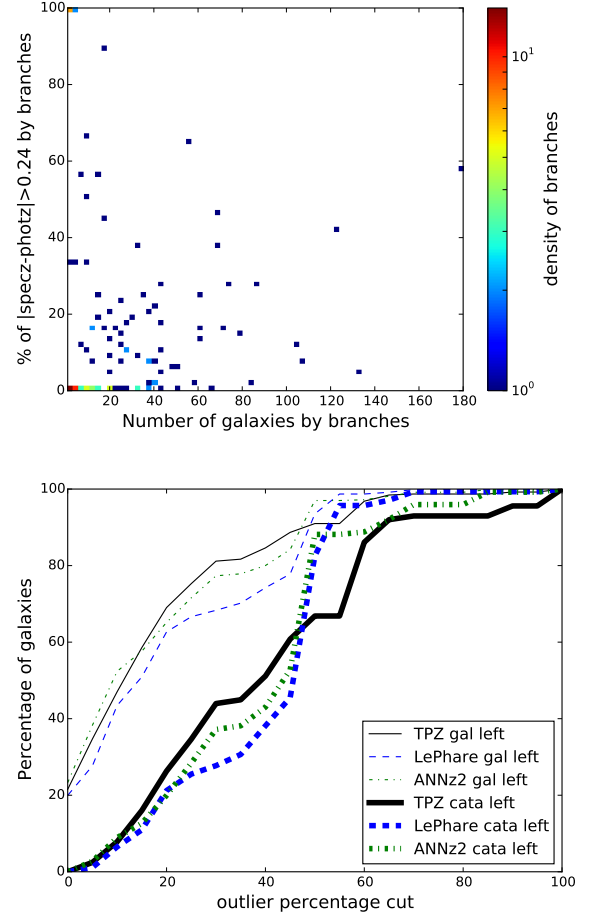


Figure 9. Top panel shows the number of galaxies as a function of the percentage of outliers by branches. Colors show the density of branches. Bottom panel shows the number of galaxies and catastrophic redshift left as a function of the outlier selection criterion applied on the branches.

- (ii) Assign the magnitude limit in each position of the galaxy catalog and in the random uniform catalog
- (iii) Using the Mangle mask, retrieve the information about the area and magnitude limit of each of the polygons into the footprint.
- (iv) Create a histogram of magnitudes limits, in our case, 10 bins from 23.4 to 25.8
- (v) Count the number of galaxies in each magnitude limit bin (magL)
- (vi) Build the density distribution in each magL as the number of galaxies over the area in the given bin and generate the density(m) function, i.e., the density as a function of magL.
- (vii) Create the Probability Function according to:

$$P[i] = \frac{\int_{m_{min}^i}^{m_{max}^i} \rho(m)}{\int_{m_{min}^{tot}}^{m_{max}^{tot}} \rho(m)}$$

where m_{max}^i and m_{min}^i are the maximum and the minimum values of magL in the bin i and m_{min}^{tot} and m_{max}^{tot} are the initial and final magL according to the binning used. In our case, $m_{max}^{tot} = 25.8$ and $m_{min}^{tot} = 23.4$.

- (viii) According to the magnitude limit where the random point lies, assign a probability to that point

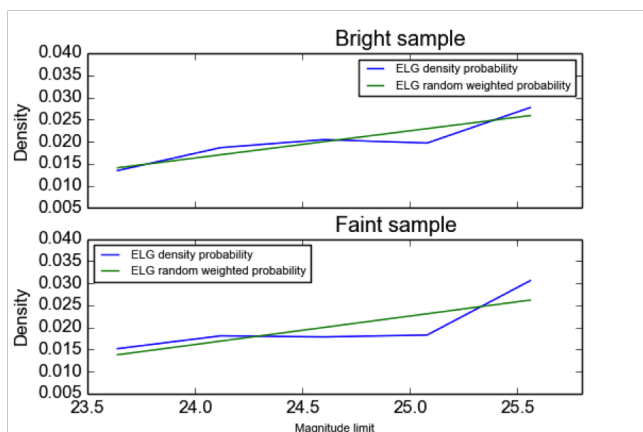


Figure 10. The blue lines show the density of targets as a function of magnitude limit in g band (given by the Mangle mask) for the bright sample (top panel) and faint sample (bottom panel). In green, the best fit to a first order polynomial, which, if we normalize to be 1 in its maximum, at $m = 25.8$, represent the weight used in the random catalog to account for the different depths of the footprint. For some intermediate magnitude limits, there is an apparent decrease in density. This is definitely a variance effect, due to the small area observed in a very inhomogeneous footprint (see figure 2) and the small number of targets.

(ix) Sample the density probability function to reject/accept random points

In Figure 10, we show the probability distribution as a function of magnitude limit according to the steps above, for both the bright and faint sample. This measurement is very limited by sample variance, nonetheless, we approximate the density distribution by a first order polynomial to assign reject/acceptance probabilities as a function of magnitude limit for the random samples. We find a mean error of 8% for each of the ELG probability densities estimated using the variance cookbook given in (Moster et al. 2011) with similar survey configurations as COSMOS.

With this probability distribution as a function of the magnitude limit in the position of the footprint, we can now calculate the random catalog used throughout the following analysis. In both samples we calculate approximately $2.8e6$ random points.

4.2 Two-point spatial correlation function

We estimate the two-point spatial correlation function using the Landy & Szalay (Landy & Szalay 1993) estimator under the fiducial cosmology over scales $1 < s < 50$ Mpc/h using the CUTE code⁷ (Alonso 2012) and calculate the galaxy bias for the samples. In this section, we use the letter s to refer to scales in redshift space for the calculation of the monopole, and r_p for scales in real space, corrected from redshift space distortions, needed to calculate the bias.

During the analysis, we have only considered Poisson errors and therefore, the uncertainty in the results presented here is underestimated. To estimate how much we are underestimating it, we calculate the cosmic variance contribution to the error in the correlation function monopole using the analytic expression given in (Xu et al. 2012). The result can be seen in figure 11. Clearly, the addition of the cosmic variance term in the error budget will worsen

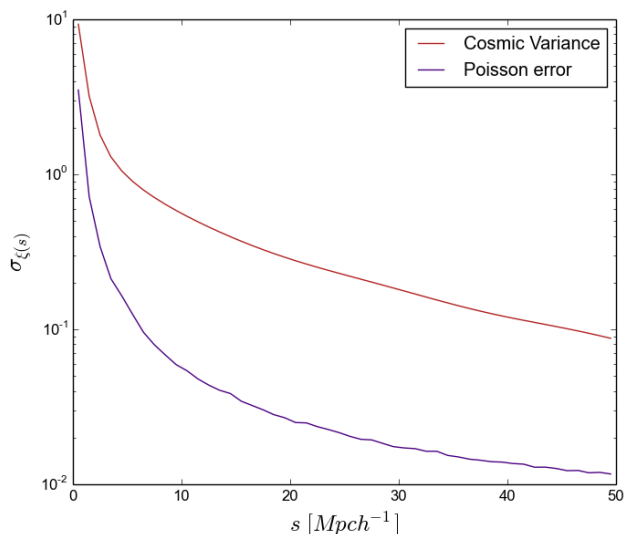


Figure 11. The cosmic variance error contribution in $\xi(s)$ compared to the Poisson unique contribution, used through the analysis. Clearly, the addition of this important term worsen the precision, so the errors given below in the analysis are underestimated compared to their true uncertainty.

the precision of our results. A more thorough analysis of the clustering signal will demand the calculation of mocks catalogs for a more precise calculation of the covariance and in a bigger volume to have a significant clustering value.

We then fit $\xi(s)$ on small scales using a power-law

$$\xi(s) = \left(\frac{s}{s_0}\right)^{-\gamma}, \quad (1)$$

In order to account for the finite volume of the sample, we measured the RR terms up to maximum separations allowed by the volume and, following (Roche & Eales 1999), we estimated an integral constraint factor (IC) as

$$IC = \frac{\sum_i \xi(s_i) RR(s_i)}{\sum_i RR(s_i)}, \quad (2)$$

where $\xi(s)$ represents the model.

We fit the power law to the data subtracting the IC from the model and allowing it to vary with the model parameters. We checked that this procedure is consistent with an iterative approach, in which we first fit a model to the original data, then use this model to estimate a correction via Eq. 2, and apply this correction to the data, repeating the process to the new data until convergence is achieved. In our case, convergence was always reached in less than 20 iterations. By fitting the model and IC correction simultaneously, we avoid the need for correcting the data.

For the fit we used comoving scales in the range $1 < s < 20h^{-1}$ Mpc. We considered different maximum scales between 10 and $50h^{-1}$ Mpc to perform the fit and found that the results are not much affected by this scale. Nonetheless $20h^{-1}$ Mpc was chosen because i) at the redshifts of interest the linear regime extends up to this scale, ii) for scales below $20h^{-1}$ Mpc the amplitude of measured $\xi(s)$ is always one order of magnitude larger than our estimates of the IC .

Our results are shown in Table 6. The z_{sp} samples are selected with spectroscopic redshifts between 0.6 and 1.2. ANNz2 and LePhare are samples selected in the same range but using z_{ph} estimates for each code; however we use their true redshifts to compute distances and $\xi(s)$.

⁷ <http://members.ift.uam-csic.es/dmonge/CUTE.html>

Sample	Redshift selection	s_0 [h^{-1} Mpc]	γ	IC	χ^2/dof
Faint	z_{sp}	$5.13^{+0.18}_{-0.19}$	$1.313^{+0.059}_{-0.062}$	0.014	0.995
	ANNz2	$5.32^{+0.19}_{-0.21}$	$1.255^{+0.060}_{-0.062}$	0.012	1.02
	LePhare	$5.30^{+0.19}_{-0.20}$	$1.265^{+0.059}_{-0.060}$	0.011	1.02
	ANNz2 LePhare-ANNz2 < 0.24	$5.57^{+0.20}_{-0.21}$	1.279 ± 0.062	0.013	1.35
	LePhare LePhare-ANNz2 < 0.24	$5.48^{+0.20}_{-0.22}$	$1.260^{+0.063}_{-0.067}$	0.013	1.15
Bright	z_{sp}	$5.27^{+0.15}_{-0.16}$	$1.264^{+0.045}_{-0.047}$	0.016	0.929
	ANNz2	$5.72^{+0.15}_{-0.16}$	$1.261^{+0.044}_{-0.047}$	0.014	1.03
	LePhare	$5.64^{+0.14}_{-0.15}$	$1.260^{+0.043}_{-0.046}$	0.014	1.12
	ANNz2 LePhare-ANNz2 < 0.24	$5.77^{+0.15}_{-0.16}$	$1.257^{+0.045}_{-0.047}$	0.016	1.06
	LePhare LePhare-ANNz2 < 0.24	$5.80^{+0.15}_{-0.17}$	$1.230^{+0.044}_{-0.046}$	0.016	1.14

Table 6. Power law fits for the monopole of the two-point correlation function of faint and bright DES target selection ELG samples. For each sample we show the impact on the parameters of the power law of selecting objects according to the different z_{ph} methods considered. $1 < s < 20$ Mpc/ h .

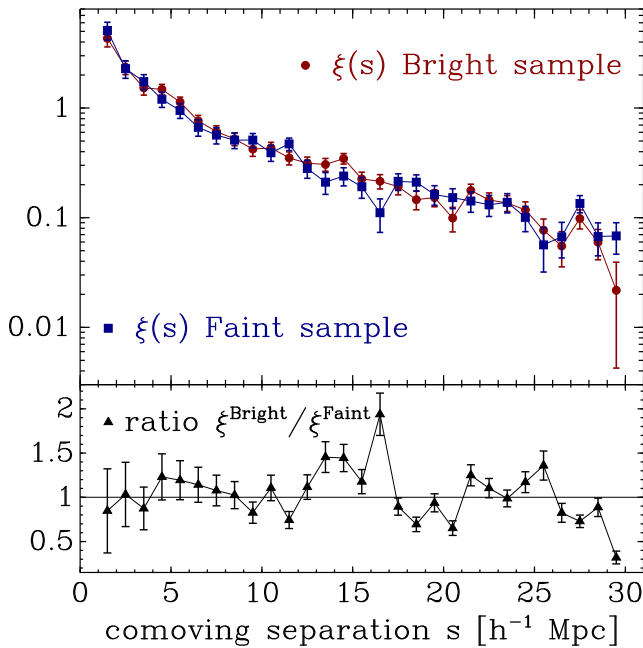


Figure 12. Comparison of the 3D two-point correlation function monopole $\xi(s)$ for the faint and bright samples. The top panel shows the measured correlations. The bottom panel displays the ratio between bright and faint samples.

A comparison of the clustering amplitudes for the bright and faint samples is shown in Figure 12. The error bars were computed propagating the uncertainties on the 2-point correlations. We see a statistical preference for the clustering amplitude of the bright sample to be higher than the one of the faint sample. This is consistent with the power law parameters fitted, Table 6, the bright sample has higher values for clustering length s_0 than the faint sample, while the slope γ seems more similar between samples. We do not draw any conclusions since we expect cosmic variance to have a large impact in these results.

Finally we find a slight increase on the clustering amplitude when photometric redshifts are considered. Figure 13 compares the amplitudes of $\xi(s)$ when z_{ph} are considered with respect to the spectroscopic selection.

4.3 Galaxy bias

In order to account for redshift space distortions, we follow the results from the VIPERS clustering analysis (Marulli et al. 2013)

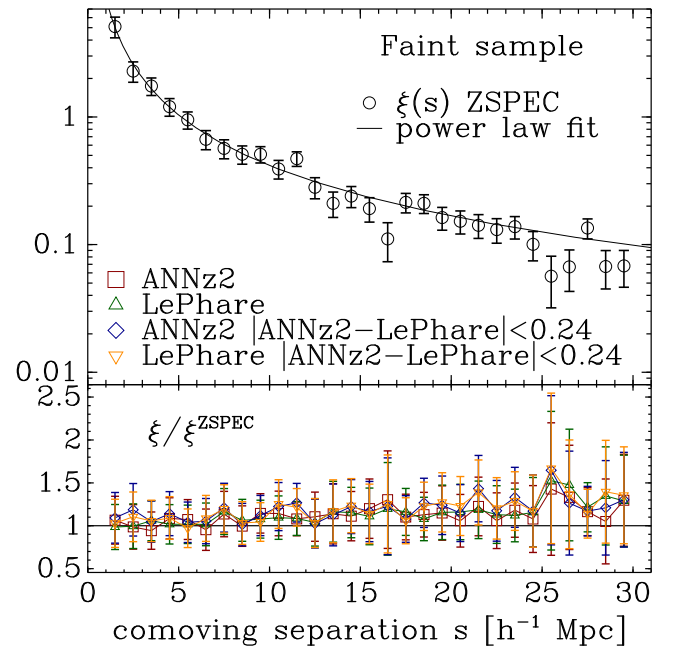


Figure 13. Comparison of $\xi(s)$ for different redshift selections for the faint sample only. The top panel shows the monopole for the clean z_{sp} sample between $0.6 < z < 1.2$. The bottom panel shows the ratio between different redshift selections and the clean sample. An apparent increase on the clustering at large scales is seen for the photometric redshifts selections.

and estimate the galaxy bias for our samples using the projected real space correlation function $w_p(r_p)$. The galaxy bias is defined as

$$b(r_p) = \sqrt{\frac{w_p(r_p)}{w_p^m(r_p)}},$$

where $w_p(r_p)$ is given by

$$w_p(r_p) = 2 \int_0^{\pi_{max}} d\pi' \xi(r_p, \pi'),$$

and is obtained from the galaxy sample, while $w_p^m(r_p)$ is the projected correlation function of matter. We compute $w_p^m(r_p)$ from the theoretical power spectrum obtained using CAMB (Lewis & Bridle 2002), with the HALOFIT routine (Smith et al. 2003) for non-linear corrections assuming a flat Λ CDM model with $\Omega_m = 0.307$ and $\sigma_8 = 0.8$.

We measure the 3D correlation function $\xi(r_p, \pi')$ in the spa-

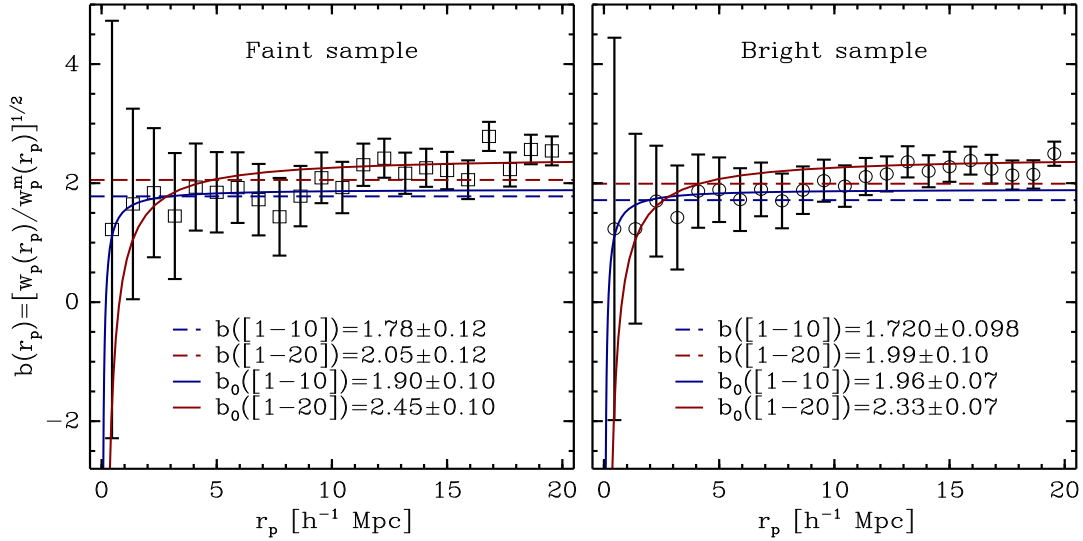


Figure 14. The galaxy bias calculated using a constant and a scale dependent relation for the faint (left) and bright (right) sample, measured from projected correlation function. In dashed lines we show the bias value calculated as the average between 1h-1Mpc and 10h-1Mpc (blue) and between 1h-1Mpc and 20h-1Mpc (red). The straight lines are the scale dependence bias, fitting to $b(r_p) = b_0 + b_1/r_p$, such that b_0 represents the large scale bias. The results depend on the range of scales used in the average and in the limiting value of the fit. For comparison to previous studies, we select the averaged bias between 1h-1Mpc and 10h-1Mpc (dashed blue) as our bias proxy, but we note that different definitions gives different results.

tial range $[1 - 40] h^{-1}\text{Mpc}$ and integrate along the line of sight to obtain the measured $w_p(r_p)$ for all samples.

For comparison with VIPERS (Marulli et al. 2013), the bias is first estimated as the average of $b(r_p)$ in the range of $[1 - 10] h^{-1}\text{Mpc}$, where the bias is fairly constant, as claimed in VIPERS and shown in Figure 14. In order to account for a small scale-dependency on the smallest scales, we also fit a relation $b(r_p) = b_0 + b_1/r_p$, such that b_0 can be taken as an estimate of the linear large-scale bias.

For both the averaged bias and the fit, the results depend on the scales used, as can be seen in Table 7. The smallest scales bring in nonlinearities whereas the largest scales are subject to sample variance, lower signal-to-noise and the largest possible effects from the *IC*. We note that we find the constant *IC* (0.01 – 0.02) to be an order of magnitude lower than the correlation (0.1 – 0.2) around $20 h^{-1}\text{Mpc}$. This 10% effect on the correlation could in principle affect the bias estimation. This effect is smaller around $10 h^{-1}\text{Mpc}$, where the correlation is a factor of 2-3 larger.

The values for the bias change significantly between the average bias and the fit to the scale dependent bias, even when the same maximum scale is used in both procedures, indicating a measurable effect of the nonlinearity on the smallest scales. Finally, the bias values for the faint and bright samples agree within the error bars for all estimations. For the linear large-scale bias definition, the bright sample bias is higher than the faint, which is counterintuitive. This is explained by the limited area of our sample (small area and big cosmic variance) and indicates that we do not have the power yet to measure robustly any galaxy bias evolution between the two target selections. Galaxy bias estimations for the faint and bright sample are within 1σ . Nonetheless, we still can compare the broad behaviour of the galaxy bias against previous measurements. The errors have been obtained propagating the uncertainties in r_0 and γ shown in the table, after fitting to (Marulli et al. 2013).

$$w(r_p) = r_p \left(\frac{r_0}{r_p} \right)^\gamma \frac{\Gamma(\frac{1}{2})\Gamma(\frac{\gamma-1}{2})}{\Gamma(\frac{\gamma}{2})}$$

In Figure 15, we compare our measurements to those published for VIPERS (Marulli et al. 2013). For this comparison, we use our averaged bias as reference to reflect the VIPERS procedure.

Our bias agrees with that from VIPERS for a population brighter than $M_B - \log(h) < -21$. To confirm this result, we calculate the absolute magnitude for the faint and bright samples together (there is a strong overlap between both samples) to directly measure the limiting absolute magnitude of our sample. We calculate the absolute magnitude for the B band using the template fitting code Lephare, fixing the redshift to its spectroscopic value. We show the B-band absolute magnitude density distribution in figure 16 as a function of redshift for the bright and faint sample. The result agrees well with what it is expected from the galaxy bias of the sample. This result should be taken with caution as we have not included the cosmic variance uncertainty. The luminosity dependent clustering will be analysed in more detail in future studies.

5 CONCLUSION

We have used 9.2 deg^2 of eBOSS test plates data to study the properties of different possible ELG target selections. We design four target selections using SDSS, SCUSS, and DES-SVA1 data. We study the SDSS-SCUSS selections in (Delubac et al. 2015; Comparat 2015) in preparation. The bright DES grz bands selection achieves 73% success rate and 71% in the desired redshift window $0.6 < z < 1.2$. The faint DES grz bands selection have slightly lower performances with 66% success rate and 68% in the redshift window. Both selections have a stellar contamination lower than 2%. We find a mean redshift of 0.8 and 0.87 for respectively the bright and faint selection.

To prepare the eBOSS survey, we looked at the strongest systematic effects which can affect the power spectrum measurement: stellar photometry contamination, airmass, galactic dust, and survey depth across the DES year one data. We find a galaxy density variation lower than 15% for each of these systematic effects,

Sample	s_0 [h^{-1} Mpc]	γ	mean z	bias averaged up to $[10 - 20] h^{-1}$ Mpc	$\chi^2/\text{d.o.f}$	b_0 fitted up to $[10 - 20] \text{Mpc} h^{-1}$
Bright	4.18 ± 0.26	1.482 ± 0.041	0.855	$[1.72 - 1.99] \pm 0.098$	2.80	$[1.96 - 2.33] \pm 0.07$
Faint	4.34 ± 0.32	1.501 ± 0.049	0.901	$[1.78 - 2.05] \pm 0.12$	3.17	$[1.90 - 2.45] \pm 0.1$

Table 7. Clustering properties and bias for the faint and bright samples selected with spectroscopic redshifts. The clustering length and slope were obtained by fitting a power law for $w(r_p)$ for $0.5 < r_p < 20h^{-1}$ Mpc. The averaged bias value was obtained by averaging the scale dependent bias $b(r_p) = [w(r_p)/w_m(r_p)]^{1/2}$, while b_0 comes from a fit to the scale dependent bias $b(r_p) = b_0 + b_1/r$. Both the average and the fit bias are obtained over scales $1 < r_p < 10h^{-1}$ Mpc as well as $1 < r_p < 20h^{-1}$ Mpc.

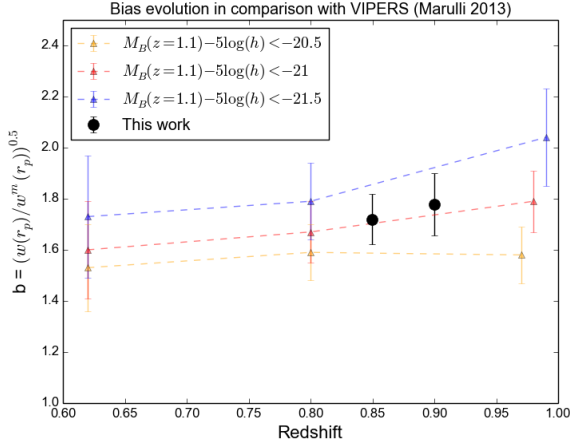


Figure 15. In black, the galaxy bias for our target selection samples in the range $0.6 < z < 1.2$ for the faint and bright. The reference values comes from the Table 1 of Marulli et al. (2013) from the VIPERS survey. In both cases biases have been measured as the average in $[1 - 10] h^{-1}$ Mpc. Our bias agrees within one sigma with a galaxy population brighter than $M_B - 5 \log(h) < -21.0$.

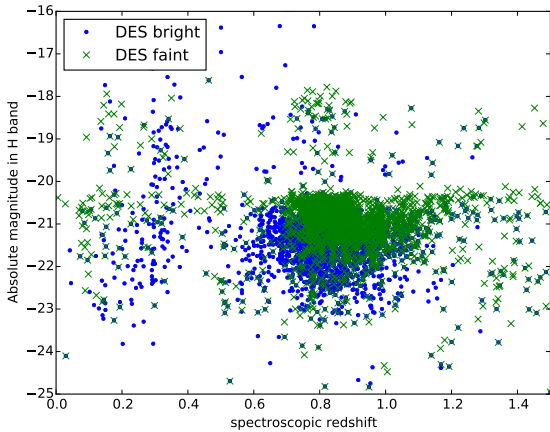


Figure 16. Absolute magnitude in the B-band for the ELG spectroscopic sample as a function of spectroscopic redshift, calculated using Lephare with the same configuration as in the photometric redshifts section. It agrees with a population $M_B - \log(h) < -21$, in the redshift interval $0.6 < z < 1.2$, as it was predicted when we measured the galaxy bias.

which is the highest fluctuation allowed to avoid damaging measurements, as studied in Dawson (2015). With a target density of 857 galaxy/deg^2 , our analysis suggests the DES bright selection will give the most accurate power spectrum measurement with an eBOSS-like survey type.

With the 4600 eBOSS z_{sp} , we investigate possible techniques to identify the z_{ph} outliers. The outlier fraction is one of the biggest source of systematics in cosmic shear and large scale structure analysis (Bernstein & Huterer 2009). With a five optical broadband photometric survey, we need to identify and control the outlier fraction. Locating color boxes with a high percentage of outliers is a possible way to deal with this. Using the random forest code TPZ, we find that removing the color branches with a percentage of outliers higher than 10%, we are left with 10% outliers and a galaxy sample of 71% completeness. Newman et al. (2013) suggests another possible technique to decrease the outlier fraction using a comparison between template fitting z_{ph} and machine learning z_{ph} . We find that in reducing the galaxy sample by 15% we decrease the outlier fraction by 30%.

We investigated the clustering properties of our samples, estimating the 3D two-point correlation function monopole $\xi(s)$ and the projected real space correlation function $w(r_p)$. We used these measurements to compute the large-scale galaxy bias, and found it to be consistent with previous ELG measurements. The galaxy bias between the DES bright and faint sample are within 1σ of each other. We find a slightly higher bias for the faint sample compared to the bright which is expected due to redshift evolution. We also looked at the binning effect in clustering analysis when having to define a redshift window with photometric redshifts. Considering that DES will have a good photometric redshifts calibration, we used spectroscopic redshifts to compute correlation functions and use the photo- z to define the 0.6 to 1.2 redshift window. We do not find significant differences when using spectroscopic and photometric redshifts. We finally compare the mean value of the galaxy bias to the deep spectroscopic survey VIPERS and find a good agreement.

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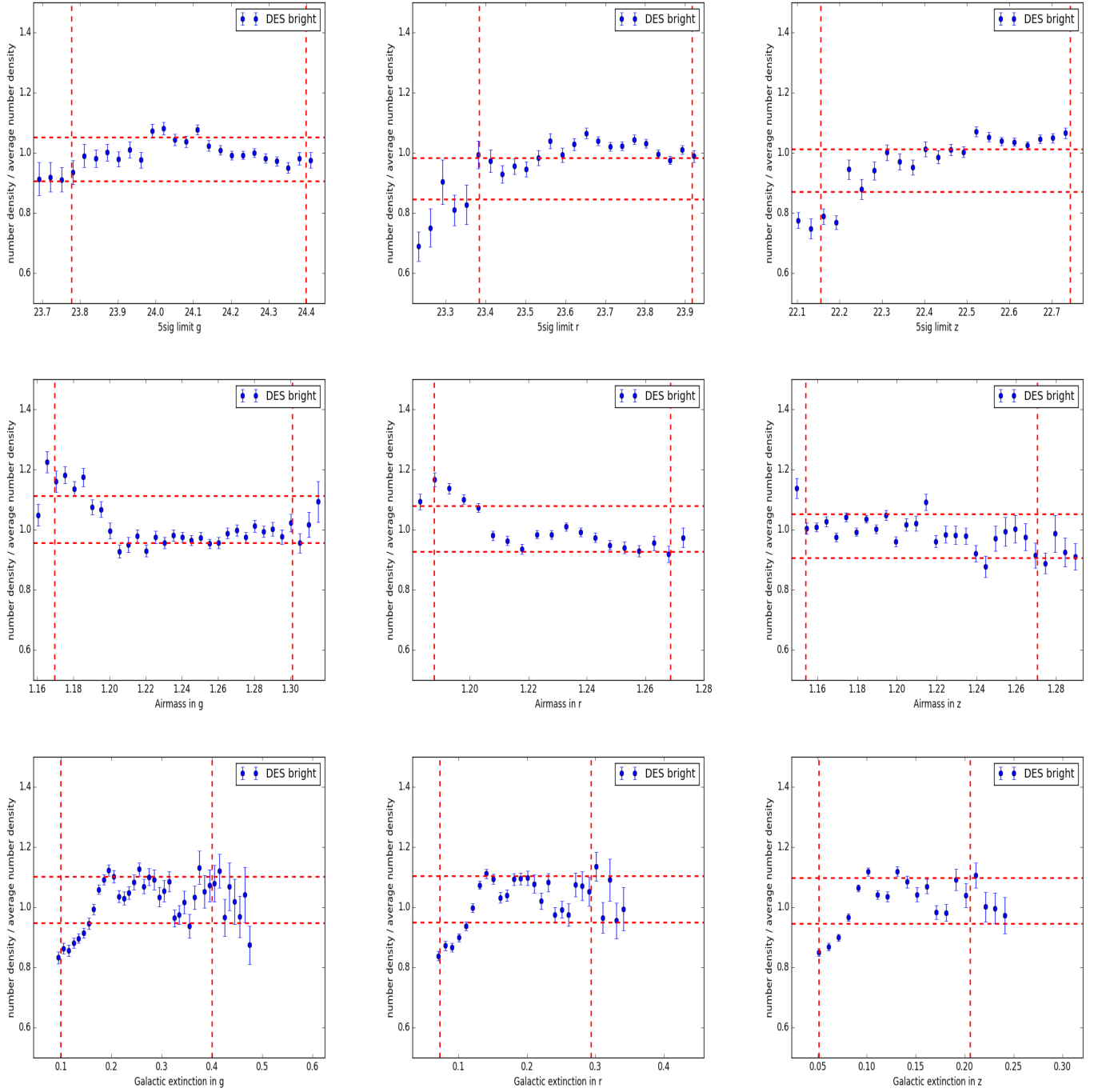


Figure A1. Density fluctuation of galaxies as a function of the depth, airmass and Galactic extinction in the top, middle and bottom row for the g,r,z bands. The two vertical and horizontal red axes show respectively the 5 and 95% of the depth, airmass, galactic extinction distribution and 15% around the mean galaxy density fluctuation for the eBOSS target selection.

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