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## DUMB-BELL GALAXIES IN SOUTHERN CLUSTERS: CATALOG AND PRELIMINARY STATISTICAL RESULTS

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### 1. Introduction

The dominant galaxy of a rich cluster is often an object whose formation and evolution is closely connected to the dynamics of the cluster itself. Hoessel (1980) and Schneider *et al.* (1983) estimate that 50 % of the dominant galaxies are either of the dumb-bell type or have companions at projected distances less than 20 kpc, which is far in excess of the number expected from chance projection (see also Rood and Leir 1979).

Presently there is no complete sample of these objects, with the exception of the listing of dumb-bell galaxies in BM type I and I-II clusters in the Abell statistical sample of Rood and Leir (1979).

Recent dynamical studies of dumb-bell galaxies in clusters (Valentijn and Casertano, 1988) still suffer from inhomogeneity of the sample. The fact that it is a mixture of optically and radio selected objects may have introduced an unknown biases, for instance if the probability of radio emission is enhanced by the presence of close companions (Stoche, 1978, Heckman *et al.* 1985, Vettolani and Gregorini 1988) a bias could be present in their velocity distribution. However this situation is bound to improve: a new sample of Abell clusters in the Southern Hemisphere has been constructed (Abell *et al.*, 1988 hereafter ACO), which has several advantages over the original northern catalog. The plate material (IIIaJ plates) is of better quality and reaches fainter magnitudes. This makes it possible to classify the cluster types with a higher degree of accuracy, as well as to fainter magnitudes. We have therefore decided to reconsider the whole problem constructing a new sample of dumb-bell galaxies homogeneously selected from the ACO survey.

### 2. Selection rules and samples

We have extracted from the ACO catalog two samples of clusters which were independently inspected on the ESO-SRC J plates by two of us (with subsamples inspected more than once at different times); the brightest member was classified according to the rules described below.

The first sample has been prepared in order to obtain a complete homogeneous volume limited sample of dumb-bell galaxies for further study of their dynamics,

radio emission etc. It has been constructed examining all the BM type I and I-II clusters, independently of any other cluster parameter.

The second sample contains all clusters, regardless of their BM type, with distance either measured or estimated (Scaramella *et al.* 1989), less than  $z = 0.07$ . This sample has been prepared in order to study the properties of dumb-bell galaxies in relation to the cluster environment (dominant/non-dominant cluster galaxy) and their space density, formation probability etc.

We classified the brightest member according to a five mark classification in which 4 denotes a dumb-bell galaxy, essentially according to the original definition given by Matthews *et al.* (1964): two objects with a difference in magnitude of at most 2 magnitudes, embedded in a common halo. In class 3 clusters there are two galaxies inside the halo, but with a large difference in luminosity (visually estimated as  $\geq 2$  magnitudes). In classes 2 and 1 there is a companion galaxy outside the envelope either with a large luminosity difference (2) or with similar luminosity but at a distance from the brightest cluster galaxy which is more than twice the radius of the latter (1). These two classes could contain a few real dumb-bells, but are mainly composed of optical pairs. Many objects will simply be satellite galaxies. Only spectroscopy and surface photometry will help to disentangle these cases. In class 0 clusters the brightest member is fairly isolated from other cluster galaxies.

A catalog of clusters of type 3 and 4 for the volume limited sample is given in Table I.

### 3. Selection effects

These selection criteria are clearly subjective and rather strongly depend on one, or a combination, of the following parameters: distance, galactic latitude, cluster richness.

*Distance* The effect of distance is to increase the number of galaxies per unit area, thus increasing the number of apparently close ellipticals. Moreover common halos are no longer clearly discernable because of their limited angular extent.

*Galactic latitude* The effect of galactic latitude is twofold: on the one hand the increasing number of stars per unit area contaminates the sample of Galaxy/Star pairs misclassified as dumb-bells beyond the magnitude at which stars and galaxies can be easily separated visually (around 17-18 mags), on the other hand increasing absorption dims the halo to the point at which a true dumb-bell is misclassified as a pair of E galaxies because of the lack of a common halo.

*Cluster Richness* The probability of chance superposition of two E galaxies is clearly dependent on the cluster richness. Especially at large distances, when the envelope is not clearly visible, this effect increases the number of interlopers.

Evidently also the plate scale and the color have important effects, but all clusters were inspected on the same plate material (ESO SERC J Sky Atlas film copies) so in this respect the present search is homogeneous.

Table I Dumb-bell galaxies in southern clusters ( $z < 0.07$ )

Name	RA (1950)	DEC (1950)	M 10	R	Class	BM type
419	3 06 00	-23 53	16.0	0	3	I-II
533	4 59 12	-22 42	15.5	0	3	III
548	5 44 54	-25 39	14.6	2	4	I-II
2401	21 55 36	-20 21	15.6	1	3	I
2412	22 01 18	-21 41	16.1	0	4	I-II
2799	0 35 06	-39 24	16.0	1	3	I-II
2800	0 35 30	-25 22	15.6	1	4	III
2819	0 43 42	-63 52	15.8	2	3	I-II
2824	0 46 06	-21 37	15.3	0	4	III
2854	0 58 36	-50 48	15.6	1	3	I-II
2860	1 01 48	-40 03	16.0	0	4	I
2911	1 23 48	-38 14	16.1	1	4	I-II
3009	2 20 18	-48 48	16.1	1	3	I
3089	3 06 12	-36 54	15.6	0	3	I-II
3093	3 09 12	-47 35	16.2	2	3	I
3094	3 09 18	-27 07	16.1	2	3	I-II
3095	3 10 18	-27 20	16.1	0	3	I-II
3104	3 12 36	-45 36	16.0	0	3	I
3111	3 16 06	-45 55	16.1	1	3	I-II
3112	3 16 12	-44 25	15.9	2	3	I
3151	3 38 24	-28 52	15.9	1	4	I-II
3158	3 41 42	-53 48	15.6	2	4	I-II
3164	3 44 42	-57 12	15.5	0	3	I-II
3225	4 08 24	-59 44	15.6	0	3	II
3266	4 30 30	-61 35	15.3	2	4	I-II
3323	5 09 24	-29 03	16.1	0	3	I
3368	5 48 24	-22 33	16.1	0	4	I-II
3390	6 23 18	-37 19	14.7	1	3	II
3391	6 25 12	-53 39	16.1	0	4	I
3395	6 26 30	-54 22	15.9	1	3	II
3497	11 57 30	-31 07	16.0	0	3	I-II
3505	12 06 06	-34 10	16.0	1	3	I-II
3528	12 51 36	-28 45	15.9	1	4	II
3530	12 52 54	-30 05	15.6	0	4	I-II
3532	12 54 36	-30 06	15.8	0	4	II-III
3556	13 21 18	-31 24	16.0	0	3	I
3559	13 27 06	-29 16	15.3	3	4	I
3560	13 29 00	-32 58	14.7	3	3	I
3570	13 43 54	-37 40	15.5	0	4	I-II
3581	14 04 36	-26 47	15.2	0	3	I
3695	20 31 36	-36 00	16.1	2	4	I
3716	20 47 54	-52 54	14.9	1	4	I-II
3744	21 04 18	-25 41	14.5	1	4	II-III
3771	21 26 06	-51 02	16.1	0	4	III
3822	21 50 36	-58 05	16.2	2	4	II-III
3880	22 25 00	-30 50	15.6	0	3	II

#### 4. Dumb-bells as brightest cluster members

Bautz-Morgan (Bautz and Morgan 1970, hereafter BM) I and I-II clusters of galaxies are dominated by an outstandingly large, luminous galaxy, almost invariably a cD, as defined by Matthews *et al.* (1964), while at the opposite extreme in BM class III clusters no member stands out against the other galaxies. Rood and Leir (1979), examining the Abell (1958) statistical sample (which is confined to clusters of richness 1 or greater) showed that 23 per cent (28 out of 112) of the dominant galaxies in BM I and I-II clusters are binary supergiant galaxies. On the contrary BM III type clusters contain only 3 per cent of multiple galaxies. This implies that the cD dominated clusters are physically different from the BM III clusters as to the multiple nature of the dominant galaxies.

The simple statistics in Table II on the present complete volume limited sample confirms the Rood and Leir (1979) results on the northern clusters sample, namely that in the clusters dominated by an outstanding bright galaxy (cD) the probability of having a dumb-bell (or a multiple component supergiant galaxy) is 2.5 times higher than in other BM type clusters.

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Table II Clusters with  $z < 0.07$

Multiplicity	I and I-II	Fm II to III
all	81	76
0 - 2	47 (58%)	64 (84%)
3 and 4	34 (42%)	12 (16%)