Supersoft X-ray sources in the Magellanic Clouds
Kahabka, P.

Published in:
Advances in Space Research

Citation for published version (APA):
SUPERSOFT X-RAY SOURCES IN THE MAGELLANIC CLOUDS

P. Kahabka*,**

* Astronomical Institute, University of Amsterdam, Kruislaan 403, NL-1098 SJ Amsterdam, The Netherlands
** Max-Planck-Institut für extraterrestrische Physik, D-85740 Garching, Germany

ABSTRACT

The Magellanic Clouds (MCs) have been mapped at soft X-rays (0.1-2.4 keV) with ROSAT during an all-sky survey and numerous pointed observations. A distinct class of supersoft X-ray sources (SSS) was discovered radiating with blackbody temperatures of a few times 10^6 K and with bolometric luminosities ~10^4 times solar. A detailed study of these sources in X-rays, UV and the optical spectral regime has just started. The nature of these sources was in debate with a black hole (Cowley et al. 1990), neutron star (Kylafis & Xilouris 1993) and white dwarf (Van den Heuvel et al. 1992) origin. The identification of SSS in the Small Magellanic Cloud with the hot central star of a planetary nebula (Wang 1991) and the hot component of a symbiotic nova (Kahabka et al. 1994) are in favor for a white dwarf (WD) origin. Van den Heuvel et al. 1992 have shown, that Cal 83 and Cal 87 in the Large Magellanic Cloud are close binaries with a WD accreting at high rates and stable burning hydrogen. Systems of a similar nature may be RXJ0513.9-6951 and 1E0035.4-7230. SSS are observed to be stable burning already for a period of ~10 years or are of transient nature or even recurrent. Such a behaviour is predicted for accreting and nuclear burning WDs if the accretion rate is within or below a critical range (Fujimoto 1982).

INTRODUCTION

From stellar evolution calculations WDs in close and wide binary systems experiencing high mass overflow rates and stable burning hydrogen have been predicted to show up as bright EUV sources (Iben 1982, Fujimoto 1982). They have been looked for in the regime of ultraviolet radiation and found in the hot component of symbiotic systems. It was natural to look for them in the close binary systems as well. But it was much more difficult to detect them there due to the faintness of the optical counterpart which is supposed to be an evolved main sequence star (Van den Heuvel et al. 1992). In these systems the accretion disk will be predominant in the optical. One had to wait for the soft X-ray surveys of Einstein and ROSAT to discover the hot component in these systems, the nuclear burning WD. Einstein found the first 4 candidates in the Clouds (Long et al. 1981, Wang et al. 1991, Wang & Wu 1992) and ROSAT added in the meantime 6 further discoveries (c.f. Fig.1, Tab.1 and references given there).

THE SAMPLE

Which systems do we observe in the MCs and what are their characteristics? They are obviously close and wide binaries and central stars of planetary nebulae. But they are also recurrent EUV (or soft X-ray) sources and transients which could be recurrent sources in case the recurrence period is quite long compared to the time these systems are observed (years). In Tab.1 we give a summary of the MC supersoft sample as it is presently known and published in the literature.
the systems listed in the table can be reconciled with nuclear burning WDs. Not all necessarily
need hydrogen-rich matter (or helium-rich matter) being supplied by accretion (e.g. the planetary
nebula nuclei) but obviously the big majority. This rises the question whether the PNe objects are
close binaries as well (c.f. Iben & Tutukov 1993, Heise et al. 1994). The optical spectra of these
systems contain valuable information about the chemical composition and the state of ionisation
of the accreted matter (exposed by the bright accretion disk, the optical companion or a nebula
if present). A clue in understanding these systems may come from the determination of the mass
of the WD, the mass of its envelope, the accretion rate onto the WD, the chemical composition
of the accreted matter and the characteristics of mass-loss due to a wind ejected from the WD
and/or the inner accretion disk in form of a jet. All these physical parameters have yet to be
determined or are in the course of being determined. We refer as one example to the work of
Heise et al. 1994 and Van Teeseling et al. 1994 where LTE model atmosphere spectral fits have
been applied to the X-ray data and WD masses have been determined. Surely information from
X-ray, UV and optical data has to be combined in order to come to a consistent description of
these laboratories. Supersoft systems have also been found in the Galaxy. Two are close binary
systems, RXJ 0925.7-4758 (Motch et al. 1994) and RX J0019+21 (Reinsch et al. 1993). A re-

DISCUSSION

As the phenomena involved with SSS are very manifold, we will restrict on a few aspects which
in the present understanding appear to be of peculiar interest. There may be other features like
possible jets present in these binary systems (c.f. Cowley et al. 1995) or the nature of the donor
star which in the future turn out to be of high interest, but the present knowledge forwards the
attention towards the hot nuclear burning WD. This has been outlined in the work of Fujimoto
1982, who set up a theory of hydrogen shell flashes on accreting WDs. One needs to supply a
minimum envelope mass (e.g. by accretion) to ignite the burning and then burn this envelope with
a constant burning rate of $\sim 1 - 4 \times 10^{-7} M_\odot \text{yr}^{-1}$ for a period dependent on the mass of the WD.
Such systems can be either short ($\sim 1 \text{yr}$) or quite long ($\sim 10^3 \text{yr}$) lived for high mass ($> 1 M_\odot$) or

![Graph](image_url)

Figure 1: ROSAT PSPC spectra of the LMC and SMC supersoft sources.
low mass (~0.5\(M_\odot\)) WDs. The long lived systems obviously are observed in the symbiotic (slow) novae and has it been claimed, that symbiotic systems contain low-mass WDs. The short-lived systems obviously just got discovered by ROSAT in the supersoft transient and recurrent sources (c.f. Fig.2) and favor high mass WDs involved. It strongly depends on the mass accretion rate what these systems look like. Van den Heuvel et al. 1992 have shown, that systems like Cal 83 and Cal 87 may experience accretion at a rate within the critical accretion rate for steady-state nuclear burning. He claims, that such accretion rates are supplied by donor stars (evolved main sequence stars) that are more massive than the WD accretors. That’s the case for the close binary systems. In the wide systems (the symbiotics) a giant is involved with high mass loss (c.f. de Kool et al. 1986, Sion and Starrfield 1994) providing such high accretion rates.

CONCLUSIONS

From evolutionary calculations (Rappaport et al. 1994) ~120 steady-state nuclear burning SSS are expected in the LMC and ~15 in the SMC. This does not yet consider the recurrent SSS which add another significant number. The fact, that only about 10 systems are observed may be related to SSS being hidden from our view due to substantial absorption in the Clouds.

Table 1: Supersoft X-ray sources - the Magellanic Clouds Sample

<table>
<thead>
<tr>
<th>Name</th>
<th>PSCP</th>
<th>kT</th>
<th>opt.</th>
<th>opt.</th>
<th>Orb.</th>
<th>Remarks</th>
<th>Ref.</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td>[e/s]</td>
<td>[eV]</td>
<td>IN</td>
<td>mag.</td>
<td>[d]</td>
<td></td>
<td></td>
</tr>
<tr>
<td>LMC</td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>RXJ0439.8-6809</td>
<td>1.38</td>
<td>18</td>
<td>-</td>
<td>B&gt;19</td>
<td>0.43</td>
<td>recurrent</td>
<td>1,15</td>
</tr>
<tr>
<td>RXJ0513.9-6951</td>
<td>0.03-3</td>
<td>30</td>
<td>-</td>
<td>HV 5682 V~17</td>
<td>4-8,15</td>
<td></td>
<td></td>
</tr>
<tr>
<td>Cal 83</td>
<td>0.98</td>
<td>40</td>
<td>+</td>
<td>V&gt;17</td>
<td>1.04</td>
<td>persistent</td>
<td>6,9,10</td>
</tr>
<tr>
<td>Cal 87</td>
<td>0.12</td>
<td>31</td>
<td>+</td>
<td>V~19</td>
<td>0.44</td>
<td>eclipsing</td>
<td>9-13</td>
</tr>
<tr>
<td>RXJ0550.0-7151</td>
<td>0.9</td>
<td>32</td>
<td>-</td>
<td></td>
<td></td>
<td></td>
<td>4,15</td>
</tr>
<tr>
<td>SMC</td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>1E0035.4-7230</td>
<td>0.38</td>
<td>41</td>
<td>-</td>
<td>V~20</td>
<td></td>
<td>eclipsing?</td>
<td>14,16</td>
</tr>
<tr>
<td>RXJ0048.4-7332</td>
<td>0.19</td>
<td>20</td>
<td>SMC3</td>
<td>V~15.5</td>
<td></td>
<td>slow nova</td>
<td>13</td>
</tr>
<tr>
<td>RXJ0058.6-7146</td>
<td>0.025</td>
<td>42</td>
<td>-</td>
<td></td>
<td></td>
<td>outburst</td>
<td>13</td>
</tr>
<tr>
<td>1E0056.8-7154</td>
<td>0.33</td>
<td>28</td>
<td>N67</td>
<td>V~16.6</td>
<td></td>
<td>PN</td>
<td>14,17</td>
</tr>
</tbody>
</table>

a) Ref.: 
Supersoft sources have been discussed as progenitors of supernovae type Ia and as systems that eventually may undergo accretion induced collapse (AIC). This topic has extensively been discussed in Livio 1994 and references therein. It is concluded, that supersoft systems having CO degenerates of mass 0.7-1.2$M_\odot$ will, if they grow due to accretion beyond the Chandrasekhar limit, experience carbon deflagration and henceforth a type Ia supernova event but are not a case for the AIC (initial WD masses in excess of 1.2$M_\odot$ are required). As type Ia supernovae comprise a rather inhomogenous class (c.f. Della Valle & Livio 1994), the progenitors may be found among the supersofts, recurrent novae but also among WD mergers. In late type galaxies CV like systems like the supersofts and recurrent novae are favored and in early type galaxies double degenerates (Della Valle & Livio 1994). It is interesting to note, that supersofts in which the WD does not grow beyond the Chandrasekhar limit, will become double degenerates (c.f. Iben & Tutukov 1984).

ACKNOWLEDGEMENT

P. Kahabka is a Human Capital and Mobility fellow.

REFERENCES