

Baseline Telescope: Design & Performance Analysis

H. Philip Stahl (h.philip.stahl@nasa.gov)

and the HabEx MSFC/JPL Design Team and the HabEx Science and Technology Definition Team

STOP Modeling indicates Baseline Telescope has ultra-stable wavefront for coronagraphy

HABITABLE EXOPLANET MISSION

HabEx is one of four missions under study for 2020 Astrophysics Decadal Survey. It will directly image and spectroscopically characterize planetary systems in the habitable zone of Sun-like stars. Additionally, HabEx will perform a broad range of general astrophysics science enabled by 115 to 2500 nm spectral range and 3 x 3 arc-minute FOV.

MECHANICAL STABILITY

Micro-thruster noise can excite telescope modes – causing LOS jitter and WFE instability. Dynamic STOP analysis determined rigid body motion between primary, secondary and tertiary mirrors. And inertial bending of primary mirror on its mount. Analysis assumes that each head has a flat 0.1 micro-Newton noise spectrum and telescope has 0.0005% critical damping.

Baseline Telescope design & performance specifications derive from science requirements

- Diameter
- Primary Mirror F/#

Diffraction Limit

- 2.5 (driven by Polarization) 400 nm
- LOS Stability (per axis on sky)
- WFE Stability
- < 0.5 mas rms < 1 to 250 pm rms

required SNR

Random Noise

16 ppt

Planet Flux Rati

Flux Ratio Noise

Contrast to FRN

550 nm. 20% BI

Reserve

Systematic Noise

Contrast Instability

40 ppt

sub-allocate this error

allocation

 $\boldsymbol{\epsilon_i} = \left(\frac{\partial \boldsymbol{\epsilon}}{\partial x_i}\right) \cdot \delta x_i$

sensitivity

22 ppt

12 ppt

Post Processing

f_pp 0.5 X

tolerance

>4 meters unobscured

Wavefront Stability is critical for internal coronagraphy.

To derive WFE stability specification:

- Start with allowable Starlight Leakage through Coronagraph,
- Calculate Contrast Leakage Sensitivities for each Zernike polynomial
- Sub-Allocate allowable Contrast Leakage between Zernike polynomials
- Then partition allowable Zernike errors between Line-of-Sight (LOS) jitter, Inertial WFE and Thermal WFE (see Conclusion).

DESIGN CONCEPT

Baseline telescope is a 4-meter off-axis unobscured TMA with a scarfed straylight baffle tube. Spacecraft surrounds telescope providing thermal isolation. For mechanical isolation, the two are connected only at a common interface ring.



Predicted rigid body motion (with 4X MUF) caused by micro-thruster noise is several orders of magnitude below LOS Jitter Specification (left) and has margin against Primary Mirror Inertial WFE Specification (right). Actual micro-thruster performance is 2X lower than specification.



THERMAL STABILITY

Zerodur® is baseline for primary mirror because its ± 5 ppb/C homogeneous CTE can be tailored to near 'zero' over operating temperature range – minimizing thermal WFE instability.





Micro-thrusters provide pointing control during science observations. Spacecraft has 4 forward thruster pods' and 4 aft pods. Forward pods have 4 heads. Aft pods have 8 heads (i.e. twice the thrust and noise).



VITH LAUNCH LOCKS SEQV = 297 PS

Primary Mirror designed to have sufficient stiffness and thermal mass to provide required WFE stability, survive launch and be easily manufactured.

PM Launch Lock system keeps launch stress < 300 psi. Baseline Coronagraph is Vector Vortex. Both Charge 4 and Charge 6 provide excellent Core Throughput and small Inner Working Angle.



Analysis of measured Zerodur mirror thermal performance indicates that Error Budget Thermal WFE allocation can be achieved if primary mirror stability is < 2 mK.

Stability is provided by primary mirror's mass, passive thermal isolation and active 'predictive' thermal control of the primary mirror and structure.

STOP Analysis indicates that science observations can proceed without interruption for anti-sun pitch of up to 45 degree and 15 degree roll (for speckle subtraction)

Plot on right shows Maximum Temperature change of 1.6 mK for a 15 degree anti-sun pitch after 50 hours 'digging' the dark hole. (45 deg pitch produces max temp change of 1.9 mK.) 15 degree roll after another 50 hours shows negligible temperature change.



Telescope Design Team

H. Philip Stahl, NASA Stefan R. Martin, JPL

Thomas Brooks, NASA J. Brent Knight, NASA

CONCLUSION STOP Analysis indicates that

				RSS Allocation	100%	50%	70%	50%	10%				
											Predict	ed Performance Margin	
	Order		r		VVC-4 Tolerance	LOS	Inertial	Thermal	Reserve	LOS	LOS	Inertial [uG]	Thermal [mK]
K	(Ν	Μ	Aberration	[pm rms]	[pm rms]	[pm rms]	[pm rms]	[pm rms]			0.02	2
				TOTAL RMS	1628.4	814	1140	814	163				
1	L	1	1	Tilt	1192.8	596.40	834.95	596.40	119.28	1	251.51	5714.58	195503.42
2	2	2	0	Power (Defocus)	1108.6	554.29	776.00	554.29	110.86	(623.99	1308.40	41266.52

Gary M. Kuan, JPL Juan Villalvazo, JPL Velibor Cormarkovic, JPL

Keith R. Warfield, JPL

NASA Aeronautics and Space Administration

Jet Propulsion Laboratory California Institute of Technology Pasadena, California

Copyright 2018. All rights reserved.

www.nasa.gov



Bijan Nemati, UAH

