

University of Groningen

## Broad iron emission line and kilohertz quasi-periodic oscillations in the neutron star system 4U 1636-53

Sanna, Andrea; Méndez, Mariano; Altamirano, Diego; Belloni, Tomaso; Hiemstra, Beike; Linares, Manuel

*Published in:*  
Monthly Notices of the Royal Astronomical Society

*DOI:*  
[10.1093/mnras/stu490](https://doi.org/10.1093/mnras/stu490)

**IMPORTANT NOTE: You are advised to consult the publisher's version (publisher's PDF) if you wish to cite from it. Please check the document version below.**

*Document Version*  
Publisher's PDF, also known as Version of record

*Publication date:*  
2014

[Link to publication in University of Groningen/UMCG research database](#)

*Citation for published version (APA):*

Sanna, A., Méndez, M., Altamirano, D., Belloni, T., Hiemstra, B., & Linares, M. (2014). Broad iron emission line and kilohertz quasi-periodic oscillations in the neutron star system 4U 1636-53. *Monthly Notices of the Royal Astronomical Society*, 440, 3275-3284. <https://doi.org/10.1093/mnras/stu490>

### Copyright

Other than for strictly personal use, it is not permitted to download or to forward/distribute the text or part of it without the consent of the author(s) and/or copyright holder(s), unless the work is under an open content license (like Creative Commons).

### Take-down policy

If you believe that this document breaches copyright please contact us providing details, and we will remove access to the work immediately and investigate your claim.

*Downloaded from the University of Groningen/UMCG research database (Pure): <http://www.rug.nl/research/portal>. For technical reasons the number of authors shown on this cover page is limited to 10 maximum.*

# Broad iron emission line and kilohertz quasi-periodic oscillations in the neutron star system 4U 1636–53

Andrea Sanna,<sup>1,2★</sup> Mariano Méndez,<sup>1</sup> Diego Altamirano,<sup>3,4</sup> Tomaso Belloni,<sup>5</sup> Beike Hiemstra<sup>1</sup> and Manuel Linares<sup>6,7</sup>

<sup>1</sup>Kapteyn Astronomical Institute, University of Groningen, PO Box 800, NL-9700 AV Groningen, the Netherlands

<sup>2</sup>Dipartimento di Fisica, Università degli Studi di Cagliari, SP Monserrato-Sestu km 0.7, I-09042 Monserrato, Italy

<sup>3</sup>Astronomical Institute, ‘Anton Pannekoek’, University of Amsterdam, Science Park 904, NL-1098 XH Amsterdam, the Netherlands

<sup>4</sup>Physics and Astronomy, University of Southampton, Southampton, Hampshire SO17 1BJ, UK

<sup>5</sup>INAF – Osservatorio Astronomico di Brera, Via E. Bianchi 46, I-23807 Merate (LC), Italy

<sup>6</sup>Instituto de Astrofísica de Canarias, Vía Láctea, E-38205 La Laguna, Tenerife, Spain

<sup>7</sup>Dept. Astrofísica, Universidad de La Laguna, E-38206 La Laguna, Tenerife, Spain

Accepted 2014 March 11. Received 2014 March 10; in original form 2014 January 21

## ABSTRACT

Both the broad iron (Fe) line and the frequency of the kilohertz quasi-periodic oscillations (kHz QPOs) in neutron star low-mass X-ray binaries (LMXBs) can potentially provide independent measures of the inner radius of the accretion disc. We use *XMM-Newton* and simultaneous *Rossi X-ray Timing Explorer* observations of the LMXB 4U 1636–53 to test this hypothesis. We study the properties of the Fe K $\alpha$  emission line as a function of the spectral state of the source and the frequency of the kHz QPOs. We find that the inner radius of the accretion disc deduced from the frequency of the upper kHz QPO varies as a function of the position of the source in the colour–colour diagram, in accordance with previous work and with the standard scenario of accretion disc geometry. On the contrary, the inner disc radius deduced from the profile of the iron line is not correlated with the spectral state of the source. The values of the inner radius inferred from kHz QPOs and iron lines, in four observations, do not lead to a consistent value of the neutron star mass, regardless of the model used to fit the iron line.

Our results suggest that either the kHz QPO or the standard relativistic Fe line interpretation does not apply for this system. Furthermore, the simultaneous detection of kHz QPOs and broad iron lines is difficult to reconcile with models in which the broadening of the iron line is due to the reprocessing of photons in an outflowing wind.

**Key words:** accretion, accretion discs – stars: individual: 4U 1636–53 – stars: neutron – X-rays: binaries.

## 1 INTRODUCTION

The energy and power density spectra (PDS) of low-mass X-ray binaries (LMXBs) change with time in a correlated way, generally following changes of the source luminosity, supporting the scenario in which these changes are a function of mass accretion rate in the system (e.g. Wijnands et al. 1997; Méndez et al. 1999; Gierliński & Done 2002). Evolution of the broad-band energy spectrum in low-luminosity systems is thought to reflect changes in the configuration of the accretion-disc flow (see review by Done, Gierliński & Kubota 2007, and references therein). Gierliński & Done (2002) find a strong correlation in the LMXB 4U 1608–52 between the position of the source in the colour–colour diagram and the truncation radius

of the inner accretion disc, which is likely driven by the average mass accretion rate through the disc. At low luminosity, the spectrum is consistent with emission from an accretion disc truncated far from the neutron star (NS); as the luminosity increases the spectrum softens and the inner radius of the accretion disc moves inwards.

In a very similar way, changes of the PDS appear to be driven by mass accretion rate. At low luminosity, when the energy spectrum of the source is hard, all timing components in the power spectrum have relatively low characteristic frequencies. These frequencies increase as the energy spectrum softens and the inferred mass accretion rate through the disc increases (see e.g. Méndez & van der Klis 1997; van der Klis 1997; Méndez et al. 1999; Méndez, van der Klis & Ford 2001; Homan et al. 2002; van Straaten et al. 2002; van Straaten, van der Klis & Méndez 2003; Altamirano et al. 2005, 2008a,b; Linares et al. 2005; van Straaten, van der Klis & Wijnands 2005; Linares, van der Klis & Wijnands 2007). The fact that fits to the energy

★E-mail: andrea.sanna@dsf.unica.it

spectra suggest that the accretion disc moves closer to the NS, and that the characteristic frequencies in the PDS increase as the luminosity increases, supports the idea that those frequencies are set by the dynamical frequencies in the accretion disc. The kilohertz quasi-periodic oscillations (kHz QPOs) are especially interesting because of the close correspondence between their frequencies and the Keplerian frequency at the inner edge of the accretion disc (e.g. Miller, Lamb & Psaltis 1998; Stella & Vietri 1998). On short time-scales (within a day or less), the frequency of the kHz QPOs increases monotonically as the source brightens, and the inferred mass accretion rate increases. However, on longer time-scales this correlation breaks down and the intensity–frequency diagram shows the so-called ‘parallel tracks’ (Méndez et al. 1999).

Broad asymmetric iron (Fe) lines have been often observed in accreting systems with the compact object spanning a large range of masses, from supermassive black holes in active galactic nuclei (AGN; see Fabian et al. 2000, for an extensive review) to stellar mass black holes (e.g. Miller et al. 2002, 2004) and NS systems (Bhattacharyya & Strohmayer 2007). The Fe  $K\alpha$  emission line at 6–7 keV is an important feature of the spectrum that emerges from the accretion disc as a result of reflection of the corona and the NS surface/boundary layer photons off the accretion disc. The mechanism responsible for the broad asymmetric profile of the line is still under discussion. Fabian et al. (1989) proposed that the line is broadened by Doppler and relativistic effects due to motion of the matter in the accretion disc. Di Salvo et al. (2005) and Bhattacharyya & Strohmayer (2007) discovered broad iron lines in the NS LMXBs 4U 1705–44 and Serpens X-1, respectively. Cackett et al. (2008) confirmed Bhattacharyya & Strohmayer (2007) results using independent observations, and also discovered broad, asymmetric, Fe  $K\alpha$  emission lines in the LMXBs 4U 1820–30 and GX 349+2. All these authors interpreted the broadening of the line as due to relativistic effects. Relativistic Fe lines have been observed at least in a dozen NS binary systems in the last decade (Di Salvo et al. 2005; Bhattacharyya & Strohmayer 2007; Cackett et al. 2008, 2009, 2010, 2012; Pandel, Kaaret & Corbel 2008; D’Ài et al. 2009, 2010; di Salvo et al. 2009; Iaria et al. 2009; Papitto et al. 2009; Reis, Fabian & Young 2009; Egron et al. 2011; Sanna et al. 2013). A different interpretation for the broadening of the line has been suggested by Ng et al. (2010) who re-analysed the data of several NS systems showing Fe  $K\alpha$  emission lines. Ng et al. (2010) claim that for most of the lines there is no need to invoke special and general relativity to explain the broad profile, and that Compton broadening is enough.

If the relativistic interpretation of the Fe line is correct, we can directly test accretion disc models by studying the line properties, as the shape of the profile depends on the inner and the outer disc radius. We expect then the iron line to be broader in the soft state – when the inner radius is smaller and the relativistic effects stronger – than in the hard state.

Accretion disc models can also be tested using simultaneous measurements of kHz QPOs and iron lines (Piraino, Santangelo & Kaaret 2000; Cackett et al. 2010). If both the Keplerian interpretation of the kHz QPOs frequency and the broadening mechanism (Doppler/relativistic) of the Fe line are correct, these two observables should provide consistent information about the accretion disc. Furthermore, if the changes in the spectral continuum also reflect changes of the inner edge of the accretion disc, the Fe line should vary in correlation with the frequency of the kHz QPOs (see e.g. Bhattacharyya & Strohmayer 2007). Understanding the relation between kHz QPOs, Fe emission line, and spectral states may have an impact beyond accretion disc physics. As discussed by Piraino

et al. (2000), Bhattacharyya & Strohmayer (2007), and Cackett et al. (2008), measurements of the line could also help constraining the mass and radius of the NS (Piraino et al. 2000), parameters needed to determine the NS equation of state. Cackett et al. (2010)<sup>1</sup> tested this idea using three observations of 4U 1636–53.

With all of this in mind, in this paper we investigate the correlation between the iron line, kHz QPOs, and spectral states in the LMXB 4U 1636–53, with the aim of understanding whether the existing interpretations of these phenomena are consistent.

The fact that 4U 1636–53 is well sampled with *XMM-Newton*, and *Rossi X-ray Timing Explorer (RXTE)* observations, is one of the most prolific sources of kHz QPOs, shows strong Fe  $K\alpha$  lines, and shows regular hard-to-soft-to-hard state transitions on time-scales of weeks, makes this source an excellent target for this study.

## 2 OBSERVATIONS

The NS LMXB 4U 1636–53 has been observed with *XMM-Newton* nine times between 2000 and 2009. The first two observations (in 2000 and 2001) were performed in imaging mode, and suffered from severe pile-up and data loss due to telemetry drop outs; we therefore excluded these two observations from the analysis. In Sanna et al. (2013) we reanalysed the 2005, 2007, and 2008 observations (Pandel et al. 2008; Cackett et al. 2010) and we analysed for the first time four new observations taken in 2009. The observation taken on 2009 March 14 had a flaring high-energy background during the full  $\sim 40$  ks exposure, and we therefore excluded this observation from the analysis (see Sanna et al. 2013, for more details on the analysis of the *XMM-Newton* observations). In this paper we make use of the results reported there.

Since 1996, 4U 1636–53 was observed  $\sim 1300$  times with *RXTE*; from 2005 March the source was regularly observed for  $\sim 2$  ks every other day (except for periods with solar viewing occultation). The first results of this monitoring campaign have been reported by Belloni et al. (2007), Zhang et al. (2013), and Sanna et al. (2012). We took all the information related to the kHz QPOs used for this paper from Sanna et al. (2012). Following Sanna et al. (2013), we refer to the 2005, 2006, 2007, 2009 March 25, 2009 September 5, and 2009 September 11 *XMM-Newton* observations as Obs. 1–6, respectively. We will use the same labelling (Obs. 1, etc.) for the simultaneous *RXTE* observations used to study the X-ray variability. When these names are used in the context of power spectral components, spectral hardness, intensities, and colours, we refer to the *RXTE* observations.

## 3 DATA ANALYSIS

### 3.1 Timing analysis

We used data from 1280 *RXTE* observations with the Proportional Counter Array (PCA; Zhang et al. 1993), covering 14 yr of data since 1996. We produced light curves and colours using the Standard 2 data as in Altamirano et al. (2008a): we used 16 s time resolution Standard 2 data to calculate the hard and soft colours, defined as the 9.7–16.0 keV/6.0–9.7 keV and 3.5–6.0 keV/2.0–3.5 keV count rate ratio, respectively. We measured the intensity in the 2.0–16.0 keV

<sup>1</sup> At about the same time, Altamirano et al. (2010) presented a similar analysis in a manuscript that was eventually never published (cf. section 5 in Cackett et al. 2010).

range. We normalized the colours and the intensity by the corresponding Crab nebula values (closest in time and within the same PCA gain epoch; e.g. Kuulkers et al. 1994; van Straaten et al. 2003).

We produced Fourier PDS using the 2–60 keV data from the  $\sim 125 \mu\text{s}$  (1/8192 s) time resolution event mode. We created Leahy-normalized PDS for each 16-s data segment with time bins of 1/8192 s, such that the lowest available frequency is 1/16 Hz, and the Nyquist frequency is 4096 Hz. We fitted the PDS in the 50–4000 Hz range using a combination of one or two Lorentzians to fit the kHz QPOs, a constant for the Poissonian noise, and if needed, an extra Lorentzian to fit the broad-band noise at frequencies between 50 Hz and the frequency range spanned by the kHz QPOs.

### 3.2 Spectral analysis

For this work we use the results of the spectral analysis by Sanna et al. (2013) that we briefly summarized hereunder.

The energy spectrum of 4U 1636–53 can be well fitted with a multicolour disc blackbody (DISKBB) plus a single-temperature blackbody (BBODY) and a thermally Comptonized component (NTHCOMP) to account for the thermal emission from the accretion disc, the thermal emission from the NS surface/boundary layer, and the high-energy emission from the corona-like region surrounding the systems, respectively. The soft seed photons of the Comptonized component were assumed to come from the accretion disc. Besides the continuum emission, the data require an emission-like feature to model prominent residuals in the energy range 4–9 keV, around the Fe  $K\alpha$  emission line region. Sanna et al. (2013) fitted this feature with a set of so-called phenomenological models (GAUSSIAN, DISKLINE, LAOR, and KYRLINE), that only model the Fe emission line, and with two reflection models (RFXCONV, and BBREFL), which self-consistently model the whole reflection spectrum, of which the Fe  $K\alpha$  emission line is the strongest feature. Sanna et al. (2013) showed that the Fe  $K\alpha$  emission line is well fitted by a symmetric Gaussian profile characterized, however, by a large breadth, which is difficult to explain with mechanisms other than relativistic broadening.

## 4 RESULTS

### 4.1 Long-term spectral behaviour of 4U 1636–53

In Fig. 1 we show the colour–colour (top), hard colour versus intensity (bottom left), and the soft colour versus intensity (bottom right) diagrams for all the available *RXTE* pointed observations. We use different symbols and colours to mark the location of the source during the *XMM-Newton* observations as estimated by the simultaneous *RXTE* observations (see Section 2).

Although 4U 1636–53 is a persistent X-ray source, it shows variations in intensity up to a factor of 6, following a narrow track in the colour–colour and colour–intensity diagrams (Belloni et al. 2007; Altamirano et al. 2008a). Using data from the All-Sky Monitor (ASM) onboard *RXTE*, Shih et al. (2005) found a long-term modulation with a period of 30–40 d, which corresponds to the regular transition between the hard and soft states (Belloni et al. 2007).

The *XMM-Newton* observations sample different parts of the diagrams shown in Fig. 1. During Obs. 1 and 6, the intensity was low and the spectrum was hard (see tables 4 and 5 in Sanna et al. 2013 for more details on the spectral properties). Obs. 2–5 were all done when the source was bright, and the source spectrum was relatively

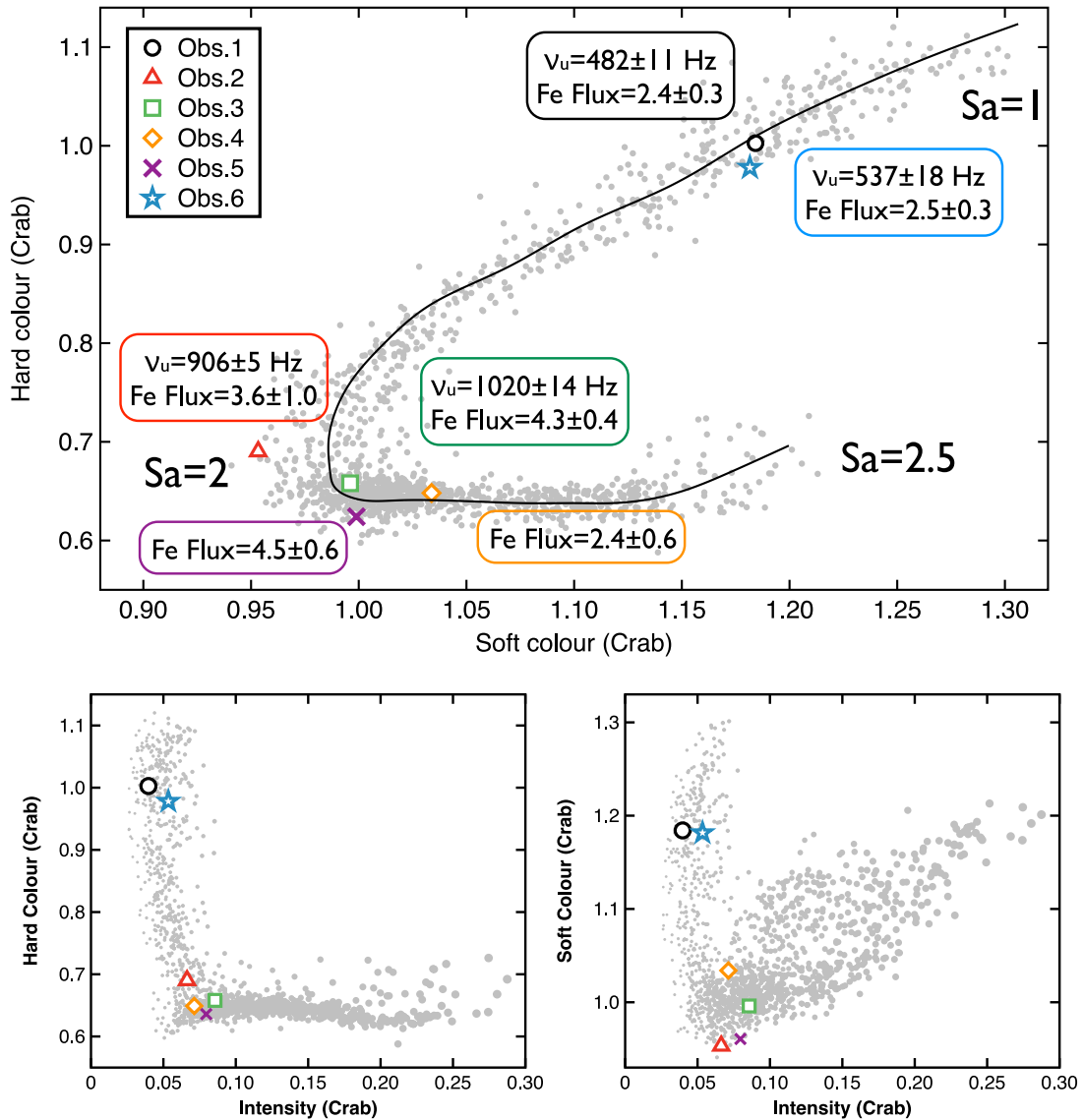
soft. The shape of the colour–colour diagram in the top panel of Fig. 1 shows that 4U 1636–53 belongs to the so-called Atoll class (Hasinger & van der Klis 1989). The upper right-hand corner of the diagram represents times when the source is in the so-called transitional state. As the mass accretion rate increases, the source first moves down to the left of the diagram, and then to right to the soft state (see e.g. van der Klis 2006). The *XMM-Newton* observations did not uniformly sample the full range of source colours and intensities, but covered two small regions of the diagram. Obs. 1 and 6 sampled the source in the transitional state, while Obs. 2–5 covered the soft state. The position of the source on the colour–colour diagram is parametrized by the length of the coordinate  $S_a$  (Méndez et al. 1999), which approximates the shape of the diagram with a curve. The length of  $S_a$  is arbitrarily normalized to the distance between  $S_a = 1$  at the top right-hand corner, and  $S_a = 2.5$  at the bottom right-hand corner, via  $S_a = 2$  at the bottom left-hand vertex of the colour–colour diagram (see Fig. 1). Similar to the  $S_z$  coordinate in the  $Z$  sources (Vrtilek, Raymond & Garcia 1990),  $S_a$  is considered to map mass accretion rate (see e.g. Hasinger & van der Klis 1989; Kuulkers et al. 1994; Méndez et al. 1999). As reported in Sanna et al. (2013), the order of the six *XMM-Newton/RXTE* observations, going from the transitional state to the soft state, according with their  $S_a$  values is 1–6–2–3–5–4.

### 4.2 Iron line and measurements of the inner accretion radius

In Table 1 we report the values of the inner disc radius for the different models used by Sanna et al. (2013) to fit the line. The observations in Table 1 are sorted according to their  $S_a$  values. As already noted in Sanna et al. (2013), the inner disc radius inferred from the relativistic profile of the iron line in 4U 1636–53 does not change in correlation with the position in the colour–colour diagram contrary to what is predicted by the standard accretion disc model (see e.g. Done et al. 2007, and references therein). Sanna et al. (2013) also found that other parameters used to fit the line, such as line energy, source inclination, and equivalent width, do not show any clear correlation with the source state (see fig. 6 in Sanna et al. 2013).

### 4.3 kHz QPOs and measurements of the inner accretion radius

We detected one or two kHz QPOs in the average PDS of each *RXTE* observation that was performed simultaneously with an *XMM-Newton* observation. In Obs. 1 and 6 we detected a strong broad-band noise component extending up to few hundred Hz plus a single kHz QPO at  $\sim 480$  and  $\sim 540$  Hz, respectively. The kHz QPOs in Obs. 1 and 6 have rms amplitudes of  $\sim 12$  and  $\sim 14$  per cent, respectively. The overall power spectral shape in both cases (not shown) is similar to those previously observed in the transitional state of 4U 1636–53 (e.g. intervals A–C in Altamirano et al. 2008a) and other sources (e.g. van Straaten et al. 2002, 2003). In Obs. 4 we detected a single kHz QPO at  $\sim 920$  Hz, with an rms amplitude of about 7 per cent. Twin kHz QPOs at  $\sim 600$  and  $\sim 905$  Hz are present in Obs. 2 with rms amplitudes of  $\sim 7$  and  $\sim 10$  per cent, respectively, and at around 700 and  $\sim 1020$  Hz in Obs. 3, both with rms amplitudes of  $\sim 6$  per cent. In Obs. 5 we detected two peaks with a frequency separation significantly lower than the average frequency difference between the lower and the upper kHz QPOs previously reported for this source (Méndez et al. 1998; Jonker, Méndez & van der Klis 2002; Méndez & Belloni 2007; Altamirano et al. 2008a). We investigated the evolution of the QPO frequency with time in



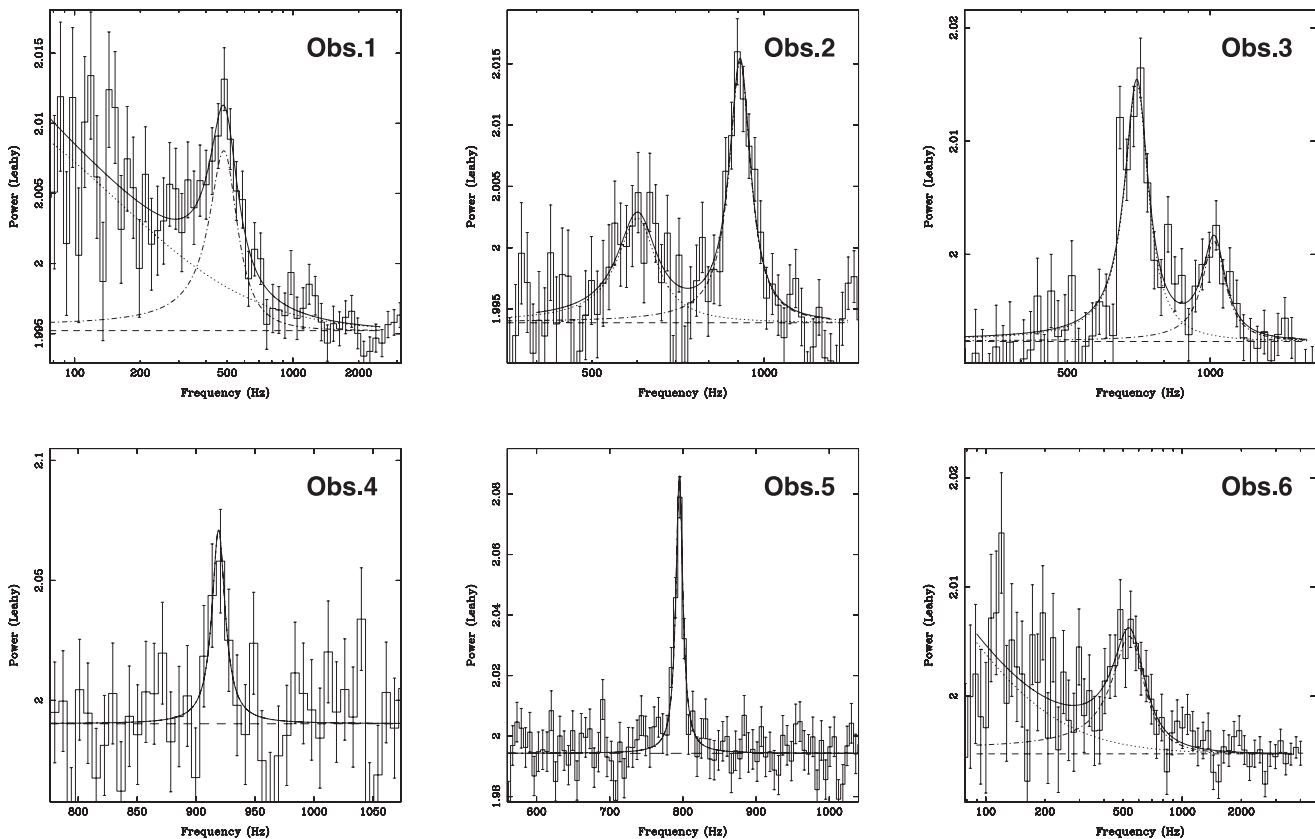
**Figure 1.** Colour–colour, hard colour versus intensity and soft colour versus intensity diagrams (upper, lower left-, and lower right-hand panels, respectively) of 4U 1636–53 for all available *RXTE* pointed observations (grey circles). Colours and intensities during *XMM-Newton* observations computed from pointed *RXTE* observations (see Section 2) are marked with different symbols following the legend at the upper left-hand corner. The position of the source on the colour–colour diagram is parametrized by the length of the black solid line that represents the coordinate  $S_a$ . For each observation we report the frequency of the upper kHz QPO (when detected) and the flux of the line (Fe flux) in units of  $10^{-3}$  photons  $\text{cm}^{-2} \text{s}^{-1}$  for fits with a KYRLINE model with  $a_* = 0.27$  (see Sanna et al. 2013, for more details).

**Table 1.** Inner radius inferred from the iron line measurements sorted as a function of the  $S_a$  parameter (from Sanna et al. 2013).

	Obs. 1 $R_{\text{in}}(GM/c^2)$	Obs. 6 $R_{\text{in}}(GM/c^2)$	Obs. 2 $R_{\text{in}}(GM/c^2)$	Obs. 3 $R_{\text{in}}(GM/c^2)$	Obs. 5 $R_{\text{in}}(GM/c^2)$	Obs. 4 $R_{\text{in}}(GM/c^2)$
DISKLINE	$10.6^{+1.5}_{-2.6}$	$8.0^{+6.3}_{-2.0*}$	$10.7^{+4.5}_{-2.4}$	$8.4^{+0.7}_{-1.5}$	$6.5^{+0.6}_{-0.5*}$	$6.0^{+2.8}_{-0.0*}$
LAOR	$10.8^{+0.6}_{-2.9}$	$6.2 \pm 1.9$	$4.0^{+5.6}_{-0.8}$	$2.3^{+0.2}_{-0.5}$	$2.0 \pm 0.3$	$2.8^{+1.2}_{-0.8}$
KYRLINE $a_* = 0$	$10.8^{+2.0}_{-1.3}$	$12.2^{+1.9}_{-2.6}$	$6.3^{+1.1}_{-0.3*}$	$13.1 \pm 1.3$	$6.2^{+0.4}_{-0.2*}$	$6.0^{+1.9}_{-0.0*}$
KYRLINE $a_* = 0.27$	$10.6^{+1.9}_{-1.2}$	$12.2^{+1.7}_{-2.6}$	$12.5^{+2.9}_{-2.0}$	$13.1 \pm 1.4$	$5.9^{+0.4}_{-0.8*}$	$5.7^{+2.0}_{-0.6*}$
KYRLINE $a_* = 1$	$9.9^{+1.9}_{-1.1}$	$12.1 \pm 2.1$	$11.8^{+3.2}_{-1.7}$	$12.9^{+3.2}_{-1.7}$	$2.6 \pm 0.1$	$5.3^{+1.4}_{-1.9}$
REFLECTION	$12.6 \pm 1.7$	$19.1^{+7.6}_{-10.8}$	$7.8^{+3.1}_{-2.7}$	$15.4 \pm 2.7$	$5.4 \pm 0.1$	$5.6^{+2.2}_{-0.3}$

Notes. Asterisk ‘\*’ means that the radius is pegged at the hard limit of the range.





**Figure 2.** Leahy normalized PDS for Obs. 1–6 of 4U 1636–53, calculated from the *RXTE* observations. The PDS were fitted with a model consisting of a constant, one or two Lorentzians to fit the kHz QPOs, and (if required) a Lorentzian to model the residual broad-band noise at low frequencies. For Obs. 5 we show the kHz QPO after we applied the shift-and-add method (see Méndez et al. 1998, for more details).

this observation, and we found that the QPO signal appears sporadically during the long observation, and when it is detectable the frequency changes with time between  $\sim 750$  and 880 Hz. It is likely that Obs. 5 showed only one kHz QPO that moved in frequency, we therefore applied the shift-and-add method (Méndez et al. 1998) to recover the properties of the QPO. Note that we did not find the upper kHz QPO (see Fig. 2), and that the frequency reported for the kHz QPO in Obs. 5 is the weighed average of the frequencies spanned by the QPO during the observation. In Fig. 2 we show the Leahy normalized PDS at high frequencies for the six observations, and in Table 2 we give the best-fitting parameters of the kHz QPOs.

In those observations in which we detected two simultaneous QPOs we can readily identify the lower and the upper kHz QPOs. For the other observations we used the frequency versus hard colour

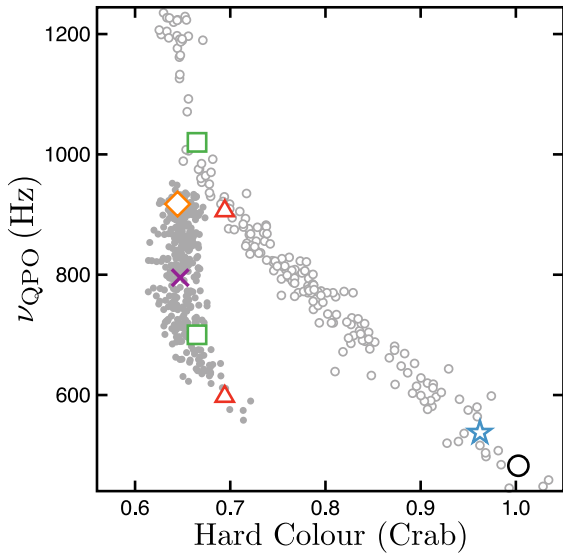
diagram that, as shown in Sanna et al. (2012, see also Belloni et al. 2007), can be used to identify the lower and upper kHz QPOs in 4U 1636–53. In Fig. 3 we show the centroid frequency of the kHz QPOs detected in 4U 1636–53 as a function of hard colour, with different symbols for the lower (grey filled bullets) and the upper (grey empty bullets) kHz QPOs. On top of that we show the kHz QPOs detected in Obs. 1–6. As expected, Fig. 3 confirms that the two simultaneous QPOs detected in Obs. 2 and 3 are indeed the lower and the upper kHz QPOs. Fig. 3 also shows that the QPOs detected in Obs. 1 and 6 are both upper kHz QPOs, while the QPOs in Obs. 4 and 5 are both the lower kHz QPOs.

In order to investigate whether the frequency of the detected kHz QPOs remains approximately constant in time, we studied the dynamical power spectra (e.g. Berger et al. 1996) using different

**Table 2.** Parameters of the kHz QPOs in 4U 1636–53 sorted as a function of the  $S_a$  parameter.

		Obs. 1	Obs. 6	Obs. 2	Obs. 3	Obs. 5	Obs. 4
$L_\ell$	$\nu$ (Hz)	–	–	$597 \pm 13$	$700 \pm 5$	$795.1 \pm 0.5$	$917 \pm 3$
	FWHM (Hz)	–	–	$115^{+42}_{-29}$	$99 \pm 14$	$11.5 \pm 1.3$	$15 \pm 6$
	rms (per cent)	–	–	$6.7 \pm 0.8$	$8.5 \pm 0.4$	$11.7 \pm 0.3$	$6.8 \pm 0.9$
$L_u$	$\nu$ (Hz)	$482 \pm 11$	$537 \pm 18$	$906 \pm 5$	$1020 \pm 14$	–	–
	FWHM (Hz)	$178 \pm 45$	$247^{+72}_{-58}$	$99 \pm 13$	$140^{+48}_{-37}$	–	–
	rms (per cent)	$11.5 \pm 1.5$	$14.1 \pm 1.4$	$9.8 \pm 0.5$	$6.3 \pm 0.7$	–	–

*Notes.*  $L_\ell$  and  $L_u$  stand for the lower and the upper kHz QPO, respectively. For Obs. 5 we report the kHz QPO parameters after applying the shift-and-add method (Méndez et al. 1998).



**Figure 3.** Frequency of the kHz QPOs in 4U 1636–53 as a function of hard colour. Grey filled and empty bullets represent the lower and the upper kHz QPOs, respectively (see Sanna et al. 2012). With large symbols we represent the kHz QPOs detected in Obs. 1–6. The symbols have the same meaning as in Fig. 1.

frequency and time binning factors. We were unable to follow the time evolution of the upper kHz QPO in any of the observations. We were able to trace the frequency of the lower kHz QPO only in two cases: During Obs. 3 the QPO frequency varied between  $\sim 620$  and  $\sim 810$  Hz while, as already mentioned, during Obs. 5 the QPO moved between  $\sim 750$  and  $\sim 880$  Hz.

In accordance with the scenario of an accretion disc truncated at a larger radii in the hard than in the soft state, and a disc extending very close to the NS surface in the soft state (Gierliński & Done 2002; Done et al. 2007), the frequency of the upper kHz QPO is lower (larger inner radius) in Obs. 6 and 1 (hard state) than in Obs. 2 and 3 (soft state).

#### 4.4 Iron lines and kHz QPOs as tracers of the inner radius of the accretion disc

Both kHz QPOs and relativistically broadened iron lines likely reflect properties of the accretion flow in the inner edge of the accretion disc (e.g. Fabian et al. 1989; Miller et al. 1998). To investigate whether the two interpretations match, we compared the inner radius estimated from the frequency of the kHz QPO and the profile of the iron line when both were detected simultaneously. Most models predict that the upper kHz QPO frequency in LMXBs represents the orbital frequency at the inner edge of the accretion disc (e.g. Miller et al. 1998; Stella & Vietri 1998). The expression for the orbital frequency in the space–time outside a slowly and uniformly rotating NS is  $\nu_\phi = \nu_k(1 + a_* (R_g/r)^{3/2})^{-1}$ , where  $\nu_k = (1/2\pi)\sqrt{GM/r^3}$  is the Keplerian frequency,  $a_* = Jc/GM^2$  is the angular momentum parameter,  $R_g = GM/c^2$  is the gravitational radius,  $G$  is the gravitational constant,  $M$  the mass of the NS,  $r$  the radial distance from the centre of the NS, and  $c$  is the speed of light.

In order to estimate the inner radius we need both the angular momentum parameter and the NS mass. As reported in Sanna et al. (2013), taking into account the spin frequency of 581 Hz of the NS in 4U 1636–53 (Zhang et al. 1997; Giles et al. 2002; Strohmayer & Markwardt 2002), we estimated  $a_*$  to be 0.27. On the other

hand, the mass of the NS in this system is unknown, therefore, we calculated the inner radius of the accretion disc,  $R_{\text{in}}$ , as a function of the NS mass. The lack of information on the NS mass does not allow us to directly compare  $R_{\text{in}}$  inferred from the kHz QPO and the iron line simultaneously, however, we can compare the trend of  $R_{\text{in}}$  as a function of other properties of the source, e.g. position in the colour–colour diagram or intensity, and test whether there is a single value of the mass of the NS for which the two different estimates of the inner disc radius agree.

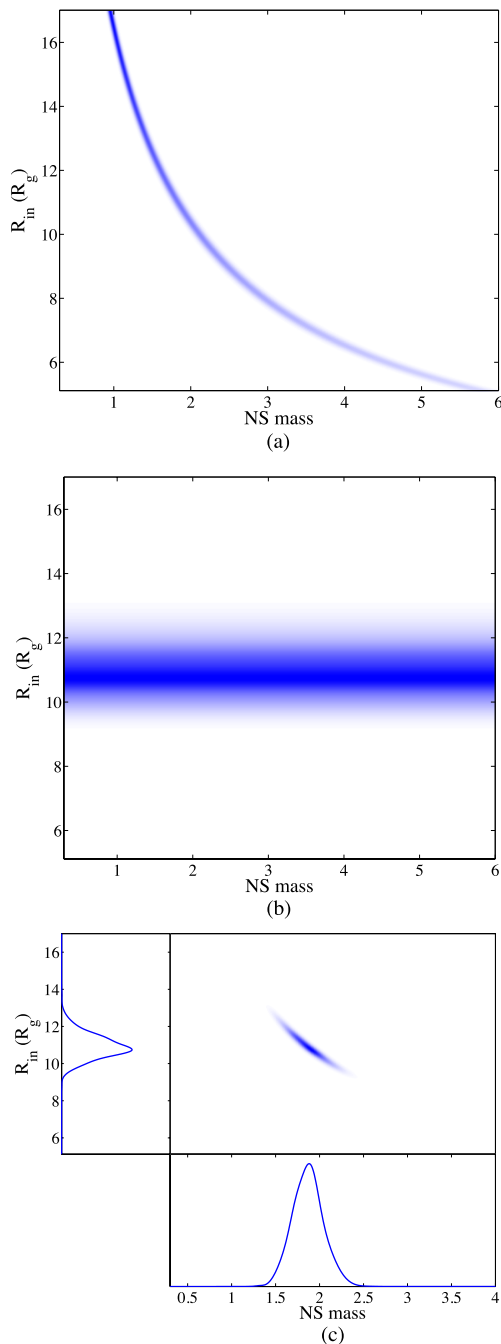
To explain in more detail how we compared the simultaneous timing and spectral information, in Fig. 4 we show a step-by-step example for the case of Obs. 1. As reported in Table 2, Obs. 1 showed an upper kHz QPO at a frequency of  $\sim 480$  Hz. In Fig. 4(a) we show the inner radius of the accretion disc in units of  $R_g$  versus the NS mass inferred from the relation between  $\nu_\phi$ ,  $M$ , and  $R_{\text{in}}$ . Colour density represents the probability distribution function (PDF) of the inner radius taking into account the uncertainties from the fit of the QPO frequency. In Fig. 4(b) we show the inner radius in units of  $R_g$  inferred from the iron emission line when we fitted it with the relativistic line model *KYRLINE* with  $a_* = 0.27$  (Sanna et al. 2013). Similar to Fig. 4(a), the PDF of the inner radius is shown in colour density. Note that the radius derived from the iron line does not depend upon the NS mass; however, for practical purposes, we plotted the PDF of the inner radius using a similar layout as for the kHz QPO. In Fig. 4(c) (central panel) we show the joint PDF of the inner radius and the NS mass derived from the iron line profile and the frequency of the kHz QPO. In Fig. 4(c) we also show the marginal distributions for the NS mass and the inner disc radius in units of  $R_g$ , calculated by integrating the joint PDF over the radius and the mass, respectively.

From this example, we find that for this observation of 4U 1636–53 the kHz QPO frequency and the iron line modelling are consistent for a NS mass of  $\sim 1.9 M_\odot$ , and an accretion disc extending down to  $\sim 11 R_g$ ; the NS mass is slightly high, but still consistent with most NS equations of state (see e.g. Lattimer & Prakash 2007). The fact that the NS mass estimate is consistent with theoretical expectations suggests that, for this particular case, the kHz QPO and iron line interpretations both hold. Since we have observations mapping different positions in the colour–colour diagram, we can test whether the kHz QPOs and the Fe line give a consistent value of the NS mass for all accretion states. To do that we used the four observations, two in the transitional state (Obs. 1 and 6) and two in the soft state (Obs. 2 and 3), where we detected both the upper kHz QPOs and the broad iron line simultaneously. For completeness, we used the inner radius estimated from fits of the iron line profile with different line models (see Table 1 for the list of models; see Sanna et al. 2013, for a discussion of the line models).

In Fig. 5 we show, for different line models, the marginal PDFs of the NS mass for the four observations with both the upper kHz QPOs and the iron emission line. The mass values derived from different observations do not yield consistent results, regardless of the iron line model. We further notice that the NS mass values in 4U 1636–53 derived from this method span a range between  $\sim 0.5$  and  $\sim 3.0 M_\odot$ , with the exception of the mass inferred from the fits of the line with the *LAOR* model, which spans a wider range ( $1.2\text{--}10 M_\odot$ ).

## 5 DISCUSSION

We detected kHz QPOs in all the *RXTE* observations simultaneous with the six *XMM-Newton* observations of the NS LMXB



**Figure 4.** (a) PDF of the inner radius of the accretion disc as a function NS mass in 4U 1636–53, inferred from the upper kHz QPO in Obs. 1. Colour density represents the confidence level. (b) PDF of the inner radius of the accretion disc as a function of NS mass in 4U 1636–53, inferred from the fit of the iron emission line in Obs. 1 using the relativistic line model KYRLINE with  $a_* = 0.27$ . Colour density represents the confidence level. (c) Joint PDF (top right-hand panel), marginal PDFs for the inner disc radius (top left-hand panel), and the NS mass (bottom panel) for 4U 1636–53, calculated combining the PDFs from (a) and (b).

4U 1636–53 for which Sanna et al. (2013) studied the broad iron line in the X-ray spectrum. Combining the measurements of the frequency of the kHz QPOs and the parameters of the iron lines in 4U 1636–53 we investigated the hypothesis that both the iron line and the kHz QPOs originate at (or very close to) the inner radius of the accretion disc in this system. From these observations we

found that the inner disc radius, deduced from the upper kHz QPO frequency, decreases as the spectrum of the source softens, and the inferred mass accretion rate increases. On the other hand, the inner radius estimated from the modelling of the relativistically broadened iron line did not show any clear correlation with the source state, except for the line model LAOR for which the inferred inner disc radius consistently decreases going from the transitional state to the soft state (see Table 1). Combining the disc radius inferred from the frequency of the upper kHz QPO and the iron line profile, we found that the mass of the NS in 4U 1636–53 deduced from the four observations is inconsistent with being the same. A similar conclusion was drawn by Cackett et al. (2010) from the first three observations in the sample that we studied here. The latter result implies that either the upper kHz QPO frequency does not reflect the orbital frequency at the inner edge of the disc, the Fe line profile is not (only) shaped by relativistic effects, the models used to fit the iron line are incorrect, or the kHz QPOs and the Fe line are not produced in the same region of the accretion disc.

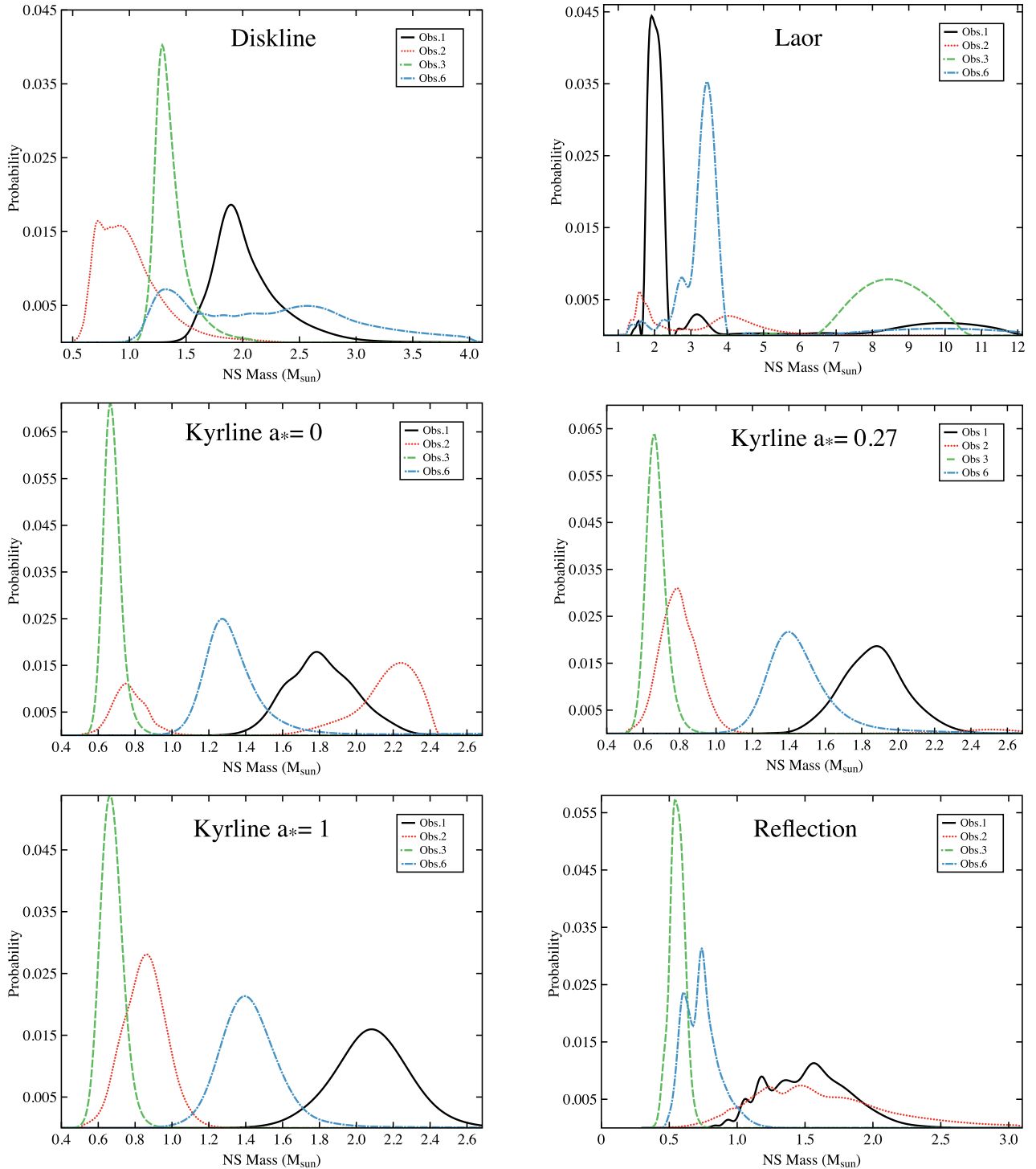
We assumed that the upper kHz QPO is the one which reflects the orbital frequency at the inner edge of the accretion disc (e.g. Miller et al. 1998; Stella & Vietri 1998). There are, however, alternative models that associate instead the lower kHz QPO to the orbital disc frequency (e.g. Meheut & Tagger 2009). If this is the case, then the radius profile showed in Fig. 4(a) would shift to higher values of  $R_{in}$ , and therefore, the mass for which  $R_{in}$  from kHz QPOs and iron lines would match will also shift to higher values. Since the difference in frequency between upper and lower kHz QPOs in 4U 1636–53 is more or less constant across the colour–colour diagram (e.g. Méndez et al. 1998; Jonker et al. 2002), using the lower kHz QPOs would lead to similar results as those shown in Fig. 5, with NS masses shifted towards higher values.

Under the assumption that the kHz QPOs are generated in the accretion disc, and considering circular orbits in the equatorial plane for Kerr space–time, the only characteristic frequencies (other than the orbital frequency) that match the observed kHz QPO frequency range are the periastron precession and the vertical epicyclic frequencies. Interpreting the upper kHz QPO as the vertical epicyclic frequency and combining the inner radius estimates with the iron line findings we found results consistent with the ones reported above. On the other hand, interpreting the upper kHz QPO as the periastron precession frequency led to meaningless NS mass values (lower than  $0.1 M_{\odot}$ ).

The kHz QPOs may still reflect the orbital (quasi-Keplerian) frequency at a radius far from the inner edge of the disc. A possible scenario to reconcile this idea, for instance, could be a mechanism that amplifies the orbital frequencies of matter orbiting within a narrow ring in the disc to produce the QPO. The process could be similar to the lamp-post model by Matt, Perola & Piro (1991). Such mechanism, however, must be able to pick a narrow range of radii in order to reproduce the observed high QPO coherence values (e.g. Barret, Olive & Miller 2005). For instance, for a  $1.8 M_{\odot}$  NS with a QPO at 800 Hz, if this is the Keplerian frequency in the disc, the putative mechanism should pick a ring of  $\sim 600$  m to produce a QPO with  $Q = 200$ . Besides generating the kHz QPOs, this mechanism should also affect other properties of the disc, such as the emissivity index or the ionization balance, which would in turn affect the properties of the iron emission line.

From the behaviour of the time derivative of the frequency of the lower kHz QPO, Sanna et al. (2012) found that the kHz QPOs (both the lower and the upper) in 4U 1636–53 are consistent with the orbital frequency at the sonic radius in the accretion disc. We also note that the frequency of the upper kHz QPO increases





**Figure 5.** Marginal PDFs of the NS mass in 4U 1636–53 inferred from simultaneous measurements of the upper kHz QPO and the iron emission line, for four different observations represented with different colours. Different panels represent different models used to fit the iron line profile.

monotonically across the colour–colour diagram. All this lends support to the interpretation of the kHz QPO reflecting the orbital frequency at the inner edge of the accretion disc.

Besides Doppler and relativistic effects, the iron emission line can be broadened by other processes. For example, the broadening may be (partially) due to Compton scattering in a disc corona (Misra & Kembhavi 1998; Misra & Sutaria 1999; see also Reynolds & Wilms 2000; Ruszkowski et al. 2000; Turner et al. 2002, and Ng et al. 2010).

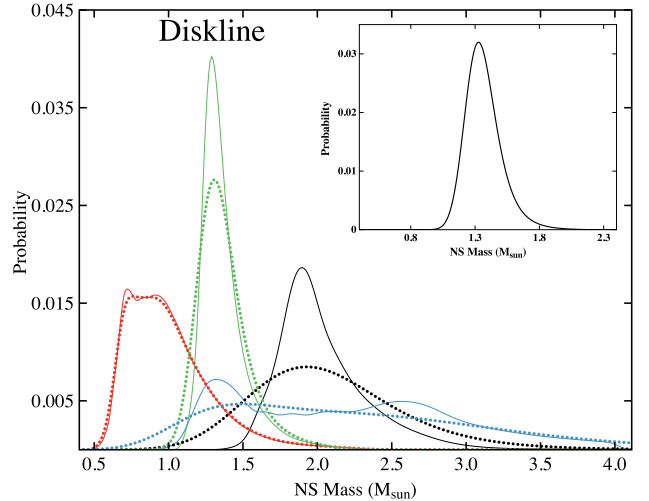
However, Sanna et al. (2013) showed that Compton broadening alone cannot explain the broad profile of the iron emission line in 4U 1636–53. Titarchuk, Kazanas & Becker (2003) argued that the red wing of the Fe  $K\alpha$  lines is not due to Doppler/relativistic effects, but to relativistic, optically thick, wide-angle (or quasi-spherical) outflows (Laming & Titarchuk 2004; Laurent & Titarchuk 2007, see however Miller et al. 2004; Miller 2007; Pandel et al. 2008). As explained by Titarchuk, Laurent & Shaposhnikov (2009), in

this scenario the red wing of the iron line is formed in a strong extended wind illuminated by the radiation emanating from the innermost part of the accreting material. One of the main predictions of this model is that all high-frequency variability should be strongly suppressed. The fact that we detected kHz QPOs and broad iron lines simultaneously in 4U 1636–53 casts doubt on this interpretation. Although our findings contradict this scenario, the model under discussion has been developed for black holes, so it is not clear how the boundary layer or the NS surface could change these predictions.

Compared to other sources, the iron line in 4U 1636–53 shows unusual properties; for instance, the best-fitting inclination is  $i \gtrsim 80^\circ$  (Pandel et al. 2008; Cackett et al. 2010; Sanna et al. 2013), which is at odds with the lack of dips or eclipse in the light curve. Pandel et al. (2008) proposed that the line profile could be the blend of two (or more) lines, for example, formed at different radii in the disc, or due to separate regions with different ionization balance. If this is correct, the total line profile would be the result of iron lines at different energies. To proceed further with this idea would require to solve the ionization balance in the accretion disc where the line is formed. Sanna et al. (2013) investigated this scenario by fitting the reflection spectrum with a self-consistent ionized reflection model, but they did not find any supporting evidence for this idea (see also Cackett et al. 2010).

The fact that in Obs. 3 and 5 the kHz QPO frequency significantly varied within the 20–30 ks required to detect the iron line suggests that the iron line profile we model may be affected by changes of the disc during those 20–30 ks. If the kHz QPO frequency depends upon the inner disc radius, the iron line profile we observe would be the average of different line profiles, one for each value of the inner disc radius. This is independent of whether the kHz QPO frequency reflects the orbital disc frequency, or whether the relation between frequency and inner radius is more complicated. The line energy or the disc emissivity could also vary if the inner disc radius changes. To proceed further, detailed simulations (assuming scenarios in which only the  $R_{\text{in}}$  changes with time, as well as scenarios in which all line parameters change) are needed to test to what extent changes in the accretion flow on time-scales of  $\sim 30$  ks (approximately the time needed with present instruments to fit the line accurately) can affect the final line profile.

A similar consideration applies to the inner radius inferred from the kHz QPO frequency. As mentioned in Section 4.3, in Obs. 3 and 5 the frequency of the lower kHz QPO spanned a frequency range of  $\sim 200$  Hz during the  $\sim 25$  ks observation. Although, we did not directly see the upper kHz QPO changing frequency with time, it is likely that the upper kHz QPO followed the lower one. If this was the case, then the full width at half-maximum (FWHM) of the upper kHz QPO observed contains information on the frequency range covered by the QPO during the observation. To bring this information into the inner disc radius estimates, we should use the QPO FWHM instead of the frequency error (which is relatively small) to calculate the PDF of the inner radius of the accretion disc. In Fig. 6 we show the marginal PDFs of the NS mass inferred from the four observations combining the upper kHz QPOs and the iron lines modelled with DISKLINE. Solid and dotted lines represent the marginal PDFs using the error in the QPO frequency and the half-width half-maximum (HWHM) as error, respectively. In the latter case the marginal PDFs of the NS mass show a broader profile, and the range of mass values where they are consistent increases (although the overlapping area is still small). By combining the marginal PDFs we get the mass profile (joint probability) for which, in the four observations, kHz QPO and iron line estimates of the inner disc radius are consistent. This is shown in the inset of Fig. 6.



**Figure 6.** Marginal PDFs of the NS mass in 4U 1636–53 using the iron line model DISKLINE. Solid and dotted lines are calculated using the error in the QPO frequency and the HWHM of the upper kHz QPOs, respectively. Colours are as in Fig. 5. The inset shows the joint PDF for the four observations with Fe line and upper kHz QPO, using the HWHM as the error in the frequency. The area under the probability function shown in the inset represents the likelihood of measuring the masses  $M_i$  of the NS mass  $M_0$  under the hypothesis  $H$  that one of the kHz QPOs is Keplerian, the Fe line is relativistic, and both phenomena arise from the same region of the accretion disc is valid.

**Table 3.** Likelihood ( $\mathcal{L}$ ) values of measuring the four values NS  $M_i$  if the NS mass  $M_0$  is always the same and under the hypothesis  $H$ , for different models of the Fe emission line. The two columns represent the likelihood measured from the marginal PDFs of the NS mass in 4U 1636–53 calculated using the error in the QPO frequency ( $\delta\nu$ ) and the HWHM of the upper kHz QPOs, respectively.

Fe line model	$\delta\nu$		HWHM	
	$M_0$ ( $M_\odot$ )	$\mathcal{L}$	$M_0$ ( $M_\odot$ )	$\mathcal{L}$
DISKLINE	$1.4 \pm 0.2$	$8.2 \times 10^{-3}$	$1.3 \pm 0.2$	$1.4 \times 10^{-1}$
LAOR	$2.6 \pm 0.6$	$< 1 \times 10^{-9}$	$3.0 \pm 1.2$	$1.8 \times 10^{-6}$
KYRLINE $a_* = 0$	$0.9 \pm 0.3$	$< 1 \times 10^{-9}$	$0.9 \pm 0.3$	$< 1 \times 10^{-9}$
KYRLINE $a_* = 0.27$	$0.8 \pm 0.2$	$< 1 \times 10^{-9}$	$0.7 \pm 0.1$	$3.1 \times 10^{-3}$
KYRLINE $a_* = 1$	$0.8 \pm 0.2$	$< 1 \times 10^{-9}$	$0.8 \pm 0.1$	$1.9 \times 10^{-3}$
REFLECTION	$0.6 \pm 0.1$	$3.7 \times 10^{-3}$	$0.6 \pm 0.1$	$2.5 \times 10^{-2}$

The most likely value of the NS mass, for this specific case, ranges between  $\sim 1.1$  and  $\sim 1.5 M_\odot$ , which is consistent with theoretical NS mass predictions (e.g. Lattimer & Prakash 2007). However, it should be noticed that in Fig. 6, two out of the four PDFs (Obs. 1 and 2) marginally overlap, therefore, the final joint probability function is likely not fully representative of all observations. Statistically speaking, the overlapping area between the intersecting marginal distributions in Fig. 6 represents the likelihood of measuring four values  $M_i$  of the NS mass  $M_0$  (assuming the mass is always the same), and the hypothesis  $H$  that one of the kHz QPOs is Keplerian, the Fe line is relativistic, and both phenomena arise from the same region of the accretion disc is valid. In Table 3 we report the NS mass values and the likelihood for the different models used to fit the Fe line profile. It is interesting to notice that the highest likelihood of the data given the model is obtained when the Fe  $K\alpha$  emission line is modelled with DISKLINE.

The frequency of the kHz QPO and the source intensity (likely the mass accretion rate) are degenerate on long (longer than  $\sim$ a day) time-scales ('parallel tracks'; Méndez et al. 1999). The same QPO frequency may appear at very different source intensities. It remains to be seen whether this phenomenon can affect some properties of the disc, such as the emissivity index or the ionization balance, which would affect the profile of the iron line, and hence the inferred value of the inner radius of the accretion disc.

## ACKNOWLEDGEMENTS

AS wish to thank Cole Miller for interesting discussions. TB acknowledges support from ASI-INAF grant I/009/10/0 and from INAF PRIN 2012-6. AS, MM, TB and DA wish to thank ISSI for their hospitality. DA acknowledges support from the Royal Society.

## REFERENCES

- Altamirano D., van der Klis M., Méndez M., Migliari S., Jonker P. G., Tiengo A., Zhang W., 2005, *ApJ*, 633, 358
- Altamirano D., van der Klis M., Méndez M., Jonker P. G., Klein-Wolt M., Lewin W. H. G., 2008a, *ApJ*, 685, 436
- Altamirano D., van der Klis M., Méndez M., Wijnands R., Markwardt C., Swank J., 2008b, *ApJ*, 687, 488
- Barret D., Olive J.-F., Miller M. C., 2005, *MNRAS*, 361, 855
- Belloni T., Homan J., Motta S., Ratti E., Méndez M., 2007, *MNRAS*, 379, 247
- Berger M. et al., 1996, *ApJ*, 469, L13
- Bhattacharyya S., Strohmayer T. E., 2007, *ApJ*, 664, L103
- Cackett E. M. et al., 2008, *ApJ*, 674, 415
- Cackett E. M., Altamirano D., Patruno A., Miller J. M., Reynolds M., Linares M., Wijnands R., 2009, *ApJ*, 694, L21
- Cackett E. M. et al., 2010, *ApJ*, 720, 205
- Cackett E. M., Miller J. M., Reis R. C., Fabian A. C., Barret D., 2012, *ApJ*, 755, 27
- D'Ai A., Iaria R., Di Salvo T., Matt G., Robba N. R., 2009, *ApJ*, 693, L1
- D'Ai A. et al., 2010, *A&A*, 516, A36
- Di Salvo T., Iaria R., Méndez M., Burderi L., Lavagetto G., Robba N. R., Stella L., van der Klis M., 2005, *ApJ*, 623, L121
- di Salvo T. et al., 2009, *MNRAS*, 398, 2022
- Done C., Gierliński M., Kubota A., 2007, *A&AR*, 15, 1
- Egron E. et al., 2011, *A&A*, 530, A99
- Fabian A. C., Rees M. J., Stella L., White N. E., 1989, *MNRAS*, 238, 729
- Fabian A. C., Iwasawa K., Reynolds C. S., Young A. J., 2000, *PASP*, 112, 1145
- Gierliński M., Done C., 2002, *MNRAS*, 337, 1373
- Giles A. B., Hill K. M., Strohmayer T. E., Cummings N., 2002, *ApJ*, 568, 279
- Hasinger G., van der Klis M., 1989, *A&A*, 225, 79
- Homan J., van der Klis M., Jonker P. G., Wijnands R., Kuulkers E., Méndez M., Lewin W. H. G., 2002, *ApJ*, 568, 878
- Iaria R., D'Ai A., di Salvo T., Robba N. R., Riggio A., Papitto A., Burderi L., 2009, *A&A*, 505, 1143
- Jonker P. G., Méndez M., van der Klis M., 2002, *MNRAS*, 336, L1
- Kuulkers E., van der Klis M., Oosterbroek T., Asai K., Dotani T., van Paradijs J., Lewin W. H. G., 1994, *A&A*, 289, 795
- Laming J. M., Titarchuk L., 2004, *ApJ*, 615, L121
- Lattimer J. M., Prakash M., 2007, *Phys. Rep.*, 442, 109
- Laurent P., Titarchuk L., 2007, *ApJ*, 656, 1056
- Linares M., van der Klis M., Altamirano D., Markwardt C. B., 2005, *ApJ*, 634, 1250
- Linares M., van der Klis M., Wijnands R., 2007, *ApJ*, 660, 595
- Matt G., Perola G. C., Piro L., 1991, *A&A*, 247, 25
- Meheut H., Tagger M., 2009, *MNRAS*, 399, 794
- Méndez M., Belloni T., 2007, *MNRAS*, 381, 790
- Méndez M., van der Klis M., 1997, *ApJ*, 479, 926
- Méndez M. et al., 1998, *ApJ*, 494, L65
- Méndez M., van der Klis M., Ford E. C., Wijnands R., van Paradijs J., 1999, *ApJ*, 511, L49
- Méndez M., van der Klis M., Ford E. C., 2001, *ApJ*, 561, 1016
- Miller J. M., 2007, *ARA&A*, 45, 441
- Miller M. C., Lamb F. K., Psaltis D., 1998, *ApJ*, 508, 791
- Miller J. M. et al., 2002, *ApJ*, 578, 348
- Miller J. M. et al., 2004, *ApJ*, 606, L131
- Misra R., Kembhavi A. K., 1998, *ApJ*, 499, 205
- Misra R., Sutaria F. K., 1999, *ApJ*, 517, 661
- Ng C., Díaz Trigo M., Cadolle Bel M., Migliari S., 2010, *A&A*, 522, A96
- Pandel D., Kaaret P., Corbel S., 2008, *ApJ*, 688, 1288
- Papitto A., Di Salvo T., D'Ai A., Iaria R., Burderi L., Riggio A., Menna M. T., Robba N. R., 2009, *A&A*, 493, L39
- Piraino S., Santangelo A., Kaaret P., 2000, *A&A*, 360, L35
- Reis R. C., Fabian A. C., Young A. J., 2009, *MNRAS*, 399, L1
- Reynolds C. S., Wilms J., 2000, *ApJ*, 533, 821
- Ruszkowski M., Fabian A. C., Ross R. R., Iwasawa K., 2000, *MNRAS*, 317, L11
- Sanna A., Méndez M., Belloni T., Altamirano D., 2012, *MNRAS*, 424, 2936
- Sanna A., Hiemstra B., Méndez M., Altamirano D., Belloni T., Linares M., 2013, *MNRAS*, 432, 1144
- Shih I. C., Bird A. J., Charles P. A., Cornelisse R., Tiramani D., 2005, *MNRAS*, 361, 602
- Stella L., Vietri M., 1998, *ApJ*, 492, L59
- Strohmayer T. E., Markwardt C. B., 2002, *ApJ*, 577, 337
- Titarchuk L., Kazanas D., Becker P. A., 2003, *ApJ*, 598, 411
- Titarchuk L., Laurent P., Shaposhnikov N., 2009, *ApJ*, 700, 1831
- Turner T. J. et al., 2002, *ApJ*, 574, L123
- van der Klis M., 1997, in Maoz D., Sternberg A., Leibowitz E. M., eds, *Astrophysics and Space Science Library*, Vol. 218, *Astronomical Time Series*. Kluwer, Dordrecht, p. 121
- van der Klis M., 2006, in Lewin W. H. G., van der Klis M., eds, *Compact Stellar X-Ray Sources*. Cambridge Univ. Press, Cambridge, in press
- van Straaten S., van der Klis M., di Salvo T., Belloni T., 2002, *ApJ*, 568, 912
- van Straaten S., van der Klis M., Méndez M., 2003, *ApJ*, 596, 1155
- van Straaten S., van der Klis M., Wijnands R., 2005, *ApJ*, 619, 455
- Vrtilek S. D., Raymond J. C., Garcia M. R., 1990, *A&A*, 235, 162
- Wijnands R. A. D., van der Klis M., van Paradijs J., Lewin W. H. G., Lamb F. K., Vaughan B., Kuulkers E., 1997, *ApJ*, 479, L141
- Zhang W., Giles A. B., Jahoda K., Soong Y., Swank J. H., Morgan E. H., 1993, in Siegmund O. H., ed., *Proc. SPIE Vol. 2006, EUV, X-Ray, and Gamma-Ray Instrumentation for Astronomy IV*. SPIE, Bellingham, p. 324
- Zhang W., Lapidus I., Swank J. H., White N. E., Titarchuk L., 1997, *IAU Circular*, 6541, 1
- Zhang G., Méndez M., Belloni T. M., Homan J., 2013, *MNRAS*, 436, 2276

This paper has been typeset from a  $\text{\TeX}/\text{\LaTeX}$  file prepared by the author.