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Discovery of two simultaneous non-harmonically related quasi-periodic oscillations in the 2005 outburst of the black hole binary GRO J1655–40

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ABSTRACT

We studied the low-frequency quasi-periodic oscillations (LFQPOs) in the black hole GRO J1655–40 during the 2005 outburst, using data from the *Rossi X-ray Timing Explorer*. All LFQPOs could be identified as either type B or type C using previously proposed classification schemes. In the soft state of the outburst the type-C LFQPOs reached frequencies that are among the highest ever seen for LFQPOs in black holes. At the peak of the outburst, in the ultraluminous state, the power spectrum showed two simultaneous, non-harmonically related peaks which we identified as a type-B and a type-C QPO. The simultaneous presence of a type-C and type-B QPO shows that at least two of the three known LFQPO types are intrinsically different and likely the result of distinct physical mechanisms. We also studied the properties of a broad peaked noise component in the power spectra of the ultraluminous state. This noise component becomes more coherent with count rate and there are strong suggestions that it evolves into a type-B QPO at the highest observed count rates.

Key words: accretion, accretion discs – binaries: close – stars: individual: GRO J1655–40 – X-rays: stars.

1 INTRODUCTION

Fast time variability is an important characteristic of black hole X-ray binaries (BHXBs) and a key ingredient in understanding the physical processes in these systems. BHXBs show a variety of X-ray spectral/variability states, representing different accretion regimes (see Belloni 2010 for a recent review). These states are often easy to identify in hardness–intensity diagrams (HID; see Homan et al. 2001). Muñoz-Darias, Motta & Belloni (2011a) recently showed that the overall strength of the rapid variability, as measured by the root-mean-square (rms) deviation, is a good tracer of these states as well. In an rms–intensity diagram (RID; see Muñoz-Darias et al. 2011a) the spectral/variability states become apparent without relying on spectral information, and the transitions between them are often better defined than in the HID. Fast (aperiodic) variability is generally studied through the inspection of power density spectra (PDS; see van der Klis 1989). Most of the power spectral components in the PDS of BHXBs are broad and can take the form of a wide power distribution over several decades of frequency or of a more localized peak (quasi-periodic oscillations, QPOs).

QPOs have been detected in many BHXBs and are thought to originate in the innermost regions of the accretion flow around the black hole. Low-frequency QPOs (LFQPOs), with frequencies ranging from a few mHz to ~ 10 Hz, are common features in BHXBs. They were first observed with *Ariel 6* in GX 339–4 (Motch et al. 1983) and observations with *Ginga* provided the first indications for the existence of multiple types of LFQPOs (see e.g. Miyamoto & Kitamoto 1991 for the case of GX 339–4 and Takizawa et al. 1997 for GS 1124–68). Observations performed with the *Rossi X-ray Timing Explorer (RXTE)* have led to an extraordinary progress in our knowledge on the properties of the LFQPOs in BHXBs (see Remillard & McClintock 2006; van der Klis 2006; Belloni, Motta & Muñoz-Darias 2011, for recent reviews). Three main types of LFQPOs are currently recognized: types A, B and C. They were originally identified in the PDS of XTE J1550–564 (see Wijnands, Homan & van der Klis 1999; Homan et al. 2001; Remillard et al. 2002), and have now been seen in several other sources (see e.g. Casella et al. 2004; Motta et al. 2011).

Despite the fact that LFQPOs have been known for several decades, their origin is still not understood, and there is no consensus about their physical nature, to the point that it is still not clear if the different types of QPOs share a common origin or

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arise from different physical phenomena.¹ However, the study of LFQPOs provides an indirect way to explore the inner accretion flow around black holes (and neutron stars). In particular, their association with specific spectral states and their phenomenology suggest that they could be a key ingredient in understanding the physical conditions that give origin to the different states.

Recently, Motta et al. (2011) showed that the differences between the three types of LFQPOs extend beyond the properties that led to the definition of the ABC classification. In particular, they showed that at least two of the three types showed a very different dependence of frequency on power-law flux, suggesting that these QPOs (types B and C) might be the result of different physical processes. Up to now different types of LFQPOs have never been detected simultaneously, hence the final proof of an intrinsic difference between different types of LFQPOs was missing. Here we present evidence of a simultaneous detection of two different types of QPOs in the PDS of the black hole binary GRO J1655–40.

1.1 GRO J1655–40

GRO J1655–40 was discovered with the *Compton Gamma-ray Observatory* during an outburst that started on 1994 July 27 (Zhang et al. 1994). Radio observations revealed superluminal jets that allowed a distance determination of 3.2 ± 0.2 kpc (Hjellming & Rupen 1995). Periodic dips have been observed in the optical and X-ray light curves during outburst (Bailyn et al. 1995; Kuulkers et al. 1998), and optical photometry and spectroscopy in quiescence provided accurate measurements of the inclination of the system ($69.5 \pm 0.1^\circ$; Orosz et al. 1997) and the mass of the black hole ($6.3 \pm 0.5 M_\odot$; Greene, Bailyn & Orosz 2001). GRO J1655–40 underwent two other major outbursts, one started in 1996 April (Remillard et al. 1996) and the other started in 2005 February (Markwardt & Swank 2005).

From variability studies of *RXTE* data spanning the 1996–1997 outburst, Remillard et al. (1999) found four types of QPOs between 0.1 and 300 Hz. Three of these QPOs had relatively stable central frequencies, whereas the central frequency of the fourth QPO varied over the range 14–28 Hz. The high-frequency QPO at 300 Hz was later confirmed by Strohmayer (2001), who also detected a QPO at 450 Hz (the highest frequency QPO seen to date from a black hole; see Belloni et al. 2011). Using the same data as Remillard et al. (1999), Sobczak et al. (2000) examined correlations between the properties of the 14–28 Hz QPO and the spectral parameters, finding that these QPOs were only detected when the hard component in the X-ray spectrum contributed more than 20 per cent to the 2–20 keV flux, and that the QPO frequency generally increased as the disc flux increased.

In 2005 *RXTE* monitored a new outburst of the source from a very early stage, observing it on daily basis and covering the full outburst phase. Saito et al. (2006) reported on the spectral evolution along the entire outburst, while Shaposhnikov et al. (2007) presented a study of the spectral and timing evolution during the early stages of the 2005 outburst. Debnath et al. (2008) performed a timing analysis to study the evolution of the variability as a function of the spectral hardening. Chakrabarti et al. (2008) studied LFQPO in

GRO J1655–40 observed during the outburst rise and decay in order to probe the origin of the frequency variations of the oscillations.

Similar to the 1996 outburst, the 2005 outburst displayed a bright ultraluminous state (ULS, also called *anomalous state*, see Belloni 2010, that in the case of GRO J1655–40 roughly corresponds to the *steep power-law state*, see Remillard et al. 1996); it lasted ~ 20 d.

2 OBSERVATIONS AND DATA ANALYSIS

We examined *RXTE*/Proportional Counter Array (PCA) public archival observations of GRO J1655–40 obtained during the 2005 outburst. For our timing analysis we only selected observations for which a narrow feature and/or a significant broad peaked component could be detected in the PDS. Our final sample includes a total of 92 observations.

The PCA data we discuss here were obtained from 2005 February through November and are now part of the public *RXTE* archive. The PCA data modes employed for most of these observations included a high time resolution event mode recording events above PCA ~ 15 keV (channel 36) and a single-bit mode covering the lower energies from ~ 2 to ~ 15 keV (channels 0–35).

As part of our analysis for each observation we computed power spectra using custom software under IDL in the energy band 2–26 keV (channels 0–62).² We used 64 s intervals and a Nyquist frequency of 1024 Hz. We then averaged the individual spectra for each observation. We did not subtract the contribution due to Poissonian noise before fitting the PDS as we experienced difficulties in applying the standard correction formula proposed by Zhang et al. (1995) and Jahoda et al. (2006). In most of the observations taken during the ULS the source luminosity is high and this results in a lower Poissonian noise level.³ The PDS were normalized according to Leahy, Elsner & Weisskopf (1983) and converted to square fractional rms (Belloni & Hasinger 1990). The integrated fractional rms⁴ was calculated over the 0.1–64 Hz frequency band. We produced a hardness–rms diagram (see Fig. 1, middle panel) following Belloni (2010) and using the method detailed in Muñoz-Darias et al. (2011a), we produced an RID (Fig. 1, bottom panel). We also examined the PDS in the form of a spectrogram (see Fig. 4), computing fast Fourier transforms of 8 s long windows of data. In some cases, we used shorter time windows to better follow the evolution of a narrow feature and to explore the evolution of the overall PDS shape over time.

PDS fitting was carried out with the standard XSPEC fitting package by using a one-to-one energy–frequency conversion and a unit response. Following Belloni, Psaltis & van der Klis (2002), we fitted the noise components with a number of broad Lorentzian shapes (one to four components), one usually zero-centred and the remaining ones centred at a few Hz. The QPOs were fitted with a variable number of Lorentzians depending on the presence of harmonic peaks. A constant component was added to all the PDS to take into account the contribution of the Poissonian noise.

STANDARD 2 mode data, with a 16 s time resolution and suitable for the spectral analysis, were used to create background and dead-time

² This choice is intended to maximize the signal-to-noise ratio of the QPO by including only the energy band where LFQPOs are significant. See Casella, Belloni & Stella (2005) and Rodriguez et al. (2004, 2008).

³ The Poissonian noise level depends on the count rate and becomes lower as the source flux increases. The Poissonian noise has been modelled for *RXTE* by Zhang et al. (1995).

⁴ We define the integrated fractional rms as the rms integrated over a certain frequency band.

¹ For theoretical models on the origin of QPOs, see e.g. Esin, McClintock & Narayan (1997), Titarchuk & Osherovich (1999), Tagger & Pellat (1999), Done, Gierlinski & Kubota (2007), Ingram & Done (2011) and references therein.

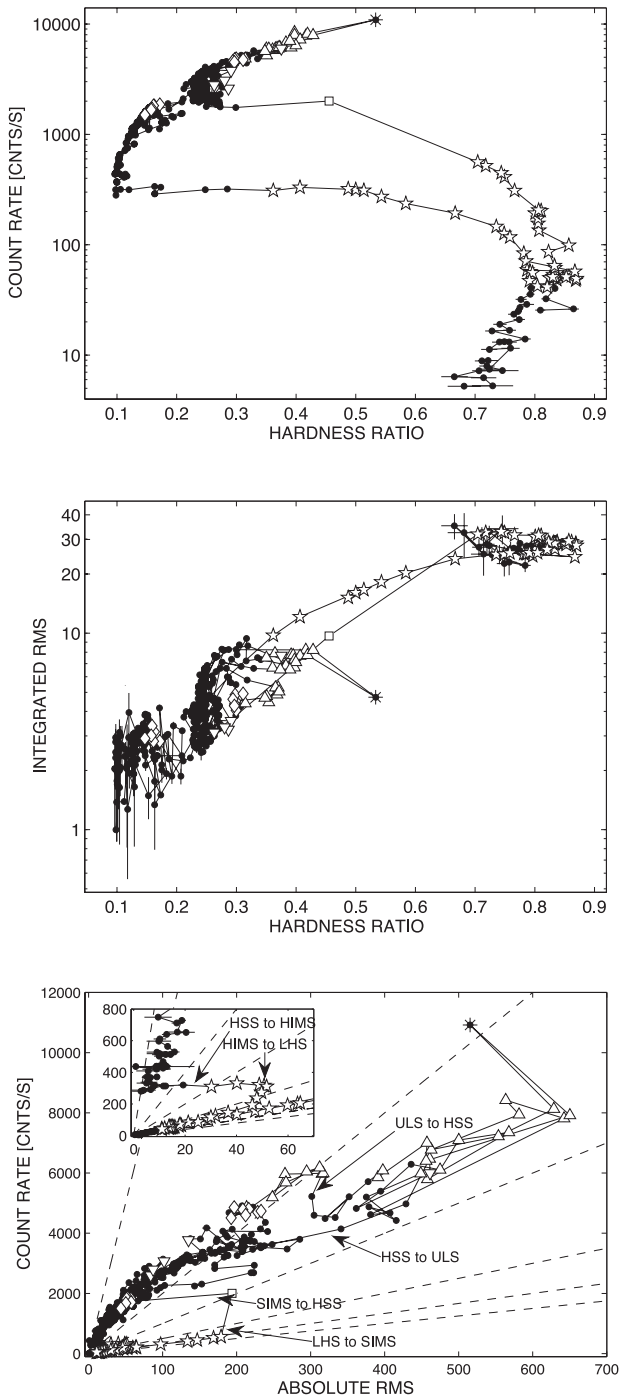


Figure 1. Upper panel: hardness–Intensity diagram. Middle panel: hardness–rms diagram. Bottom panel: rms–intensity diagram. The dashed lines represent (from left to right) the 1, 5, 10, 20, 30, 40 per cent fractional rms levels. The inset shows a close-up view of the hardening phase of the outburst. Arrows indicate the state transitions described in Section 3. Notice that the transitions to and from the ULS have been identified through a detailed timing analysis and through the inspection of the HID/RID. The source starts its outburst in the bottom left-hand corner and follows the diagram in an anticlockwise way. For further details on the RID, see Muñoz-Darias et al. (2011a). In all the panels the solid line joins observations as a function of the time. Different symbols mark the PDS type described in the text (see Section 3.1.1) and shown in Fig. 3. Stars: type 1; upward triangles: type 2; diamonds: type 3; downward triangles: type 4; square: type 5; asterisk: type 6. Black circles mark observations where no QPO was observed. See Section 2 for information about the diagrams.

corrected spectra. We extracted energy spectra for each observation using the standard *RXTE* software within *HEASOFT*, v. 6.12. Only data coming from Proportional Counter Unit 2 (PCU2) of the PCA were used for the analysis, as it was the only unit that was active during all observations. A systematic error of 0.6 per cent was added to the PCU2 spectra to account for calibration uncertainties.⁵ We accumulated background corrected PCU2 rates in the *STANDARD* 2 channel bands $A = 4–44$ (3.3–20.2 keV), $B = 4–10$ (3.3–6.1 keV) and $C = 11–20$ (6.1–10.2 keV) to produce the HID shown in Fig. 1 (upper panel). The hardness is defined as $H = C/B$ (Homan & Belloni 2005).

3 RESULTS

In this section we first describe the general behaviour of GRO J1655–40 during the full outburst. Then we focus on the timing analysis of the observations showing QPOs and/or broad peaked components.

3.1 Outburst evolution and the ultraluminous state

The HID in Fig. 1 (upper panel) shows a rather typical outburst evolution, with the addition of a complex ULS (the upper diagonal extension, above 2000 counts s^{-1}). For a comparison with HIDs of other BHXBs we refer to Dunn et al. (2010) and Fender, Homan & Belloni (2009). The ULS is also clearly visible in the RID (the long extension between 5 and 10 per cent rms levels crossing the diagram from the left- to the upper right-hand corner of the plot). The spectral classification of the observations in our sample was performed following the criteria outlined in Belloni et al. (2011). In addition to the ULS, we identified four different spectral states: low-hard state (LHS), hard-intermediate state (HIMS), soft-intermediate state (SIMS) and high-soft state (HSS).

For a period of about a month after the beginning of the *RXTE* daily monitoring, GRO J1655–40 remained in the LHS (see Fig. 2). After this the source made an extremely fast transition (in about 1 d) to the SIMS. Even though it was expected, the HIMS was not sampled by *RXTE* during this transition, while the subsequent SIMS lasted only 1 d. Joinet, Kalemci & Senziani (2008) reported the presence of a short HIMS (but no SIMS) using data from the *INTEGRAL* satellite. After the SIMS GRO J1655–40 reached the HSS, where it remained for about 2 months, with minor spectral and luminosity changes. This was followed by a period of ~ 20 d, during which the source was in the ULS. Intense flaring on time-scales of hours to days and significant spectral changes were observed during this period (see Fig. 2). The count rate also reached its maximum in the ULS ($\sim 11\,150$ counts s^{-1} PCU $^{-1}$ in the 2–26 keV energy range). After the ULS the source made a transition to the HSS; the count rate decreased continuously for ~ 20 d, while the hardness ratio remained almost constant. The duration of the HSS was 4 months. Differently from what happened during the softening phase (upper horizontal branch in the HID, 1, upper panel), *RXTE* was able to catch the short-lived (few days) HIMS during the hardening phase (bottom horizontal branch in the HID). Finally, the source moved to the LHS (arrow in the inset of Fig. 1, bottom panel) and reached quiescence after about 2 months.

The exact position of the transitions from and to LHS can be easily identified through the RID. They form sharp breaks between

⁵ See <http://www.universe.nasa.gov/xrays/programs/rxte/pca/doc/rmf/pcarmf-11.7> for a detailed discussion on the PCA calibration issues.

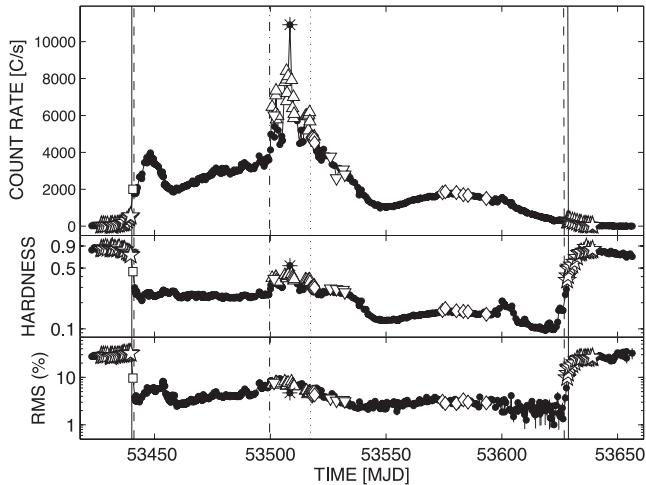


Figure 2. Upper panel: light curve of the 2005 outburst of GRO J1655–40. Middle panel: evolution of the spectral hardness. Bottom panel: evolution of the integrated fractional rms. The vertical lines mark the state transitions. First solid line: LHS to SIMS. First dashed line: SIMS to HSS. Dot–dashed line: HSS to ULS. Dotted line: ULS to HSS. Second dashed line: HSS to HIMS. Second solid line: HIMS to LHS. The symbols are the same as in Fig. 1.

the branches traced out during the HIMS and the LHS (see Fig. 1, bottom panel). In the HID these transitions appear smoother.

3.1.1 Evolution of the timing properties

To study the fast time-variability properties of GRO J1655–40, we divided the PDS in our sample (i.e. only those showing QPOs or peaked components) into different groups based on the type A/B/C classification (Wijnands et al. 1999; Casella et al. 2005; Motta et al. 2011). All the narrow features could be classified as one of the three types of QPOs, even though in a few cases the classification remains ambiguous. Our PDS and QPO classifications are detailed in Table 2.

Among the 92 observations we analysed, we identify six different types of PDS. In Fig. 3 we show an example for each type. The reader should notice that the following types of PDS do not exactly correspond to the ones described in Belloni (2010), even though some correspondences can be found (see below).

(i) *Type 1.* The PDS show two main components: a strong flat-topped noise – that can be fitted by three broad Lorentzians – and one or more QPO peaks. When more than one QPO peak is observed, they are harmonically related. The fundamental peak moves in the 0.1–13 Hz frequency range (0.1–3 Hz during the rise phase of the outburst and 0.2–13 Hz during the decay) and is strong and narrow (*quality factor*⁶ $Q > 4$). This type of PDS corresponds to PDS 1 and 2 in Belloni (2010), section 3.3.1. The integrated fractional rms of the PDS is always higher than 10 per cent. These PDS are observed in the LHS during the outburst rise and both during the HIMS and LHS in the decay.

(ii) *Type 2.* The PDS show a broad noise component that can be described by three broad Lorentzians, a narrow QPO peak at

~ 20 Hz and a peaked noise component with a characteristic frequency between 6 and 8 Hz. The broad peaked component appears in the PDS in correspondence to local peaks in the count rate. The light curve shows significant flaring also on short time-scales (hundreds of seconds). These PDS show typical fractional rms in the range 4–10 per cent and they are observed during the ULS.

(iii) *Type 3.* A weak (but significantly detected) and narrow QPO appears at frequencies between 20 and 28 Hz and is associated with a weak flat-topped noise component with a break at few Hz plus a stronger power-law noise component dominating at lower frequencies. These PDS are found in the HSS, some close to the transition between ULS and HSS and others in a more typical (softer) HSS. The QPOs have lower frequencies (19–21 Hz) and appear weak and broad ($Q \sim 3$ –10) near the ULS/HSS transitions, while they show higher frequencies (25–28 Hz) and are strong and narrow ($Q \sim 15$ –39) in softer parts of the HSS. Averaging the PDS of the two subgroups we find a peak at 20.1 Hz for the first group and 27.2 Hz for the second (soft HSS) group, with $Q = 4.4$ and 16.6, respectively. In both cases the break associated with the noise appears sharper in the averaged PDS than in the individual ones. This PDS corresponds to PDS 5 in Belloni (2010).

(iv) *Type 4.* The PDS show a peaked noise component (slightly weaker and broader with respect to the PDS of type 2) with no significant QPO. Differently from PDS type 2, the type 4 PDS do not appear associated with local maxima in the light curve, even though the light curve shows moderate flaring on short time-scales (hundreds of seconds). These PDS are mostly found in the HSS close to the transition from ULS and HSS. One case (Obs #30) is observed in the ULS.

(v) *Type 5.* This type of PDS is found within Obs #28, which consists of two *RXTE* orbits. The inspection of the spectrogram from the entire observation revealed that the shape of the PDS is variable during the observation. During the first orbit (first section of the observation) the light curve shows low variability and the PDS reveal a strong flat-topped noise component and no significant QPO. In the first half of the second orbit (second section, ~ 4900 – 5900 s after the observation start time) the PDS show a weaker power-law noise component and two harmonically related peaks, at ~ 3 and ~ 6 Hz. In the second half of the second orbit (third section), the PDS show power-law noise and a strong peak at ~ 6 Hz with a faint subharmonic at ~ 3 Hz. The first section of Obs #28 results in a PDS shape very similar to type 1 PDS (even though no QPO is detected), while the second and third section of Obs #28 result in the PDS shape that we classify as type 5 (see Fig. 3). During the entire second orbit, the light curve shows dips⁷ 4–8 s long, during which the PDS take the form of strong flat top noise with no QPO. In Fig. 3 we only show the PDS extracted from the second part of the second orbit, where the QPO centroid peak is stronger. This PDS is observed in the SIMS and corresponds to PDS 3 in Belloni (2010).

(vi) *Type 6.* Our sample contains a very peculiar PDS (Obs #42) that shows two non-harmonically related peaks at 6.84 ± 0.03 and 18.7 ± 0.1 Hz superimposed on strong power-law noise. These two simultaneous non-harmonically related peaks were first reported by Homan et al. (2005b). A careful inspection of the spectrogram confirms that the two peaks are always present during the full length of the observation. The peak at ~ 6 Hz has fractional rms = 0.81

⁶ The *quality factor* Q is defined as the ratio of the frequency of a Lorentzian and its full width at half-maximum (FWHM). Conventionally, a QPO is defined as a feature with a quality factor larger than 2.

⁷ We studied the properties of those dips and we verified the absence of hardening, therefore, we exclude an absorption origin (see Kuulkers et al. 1998).

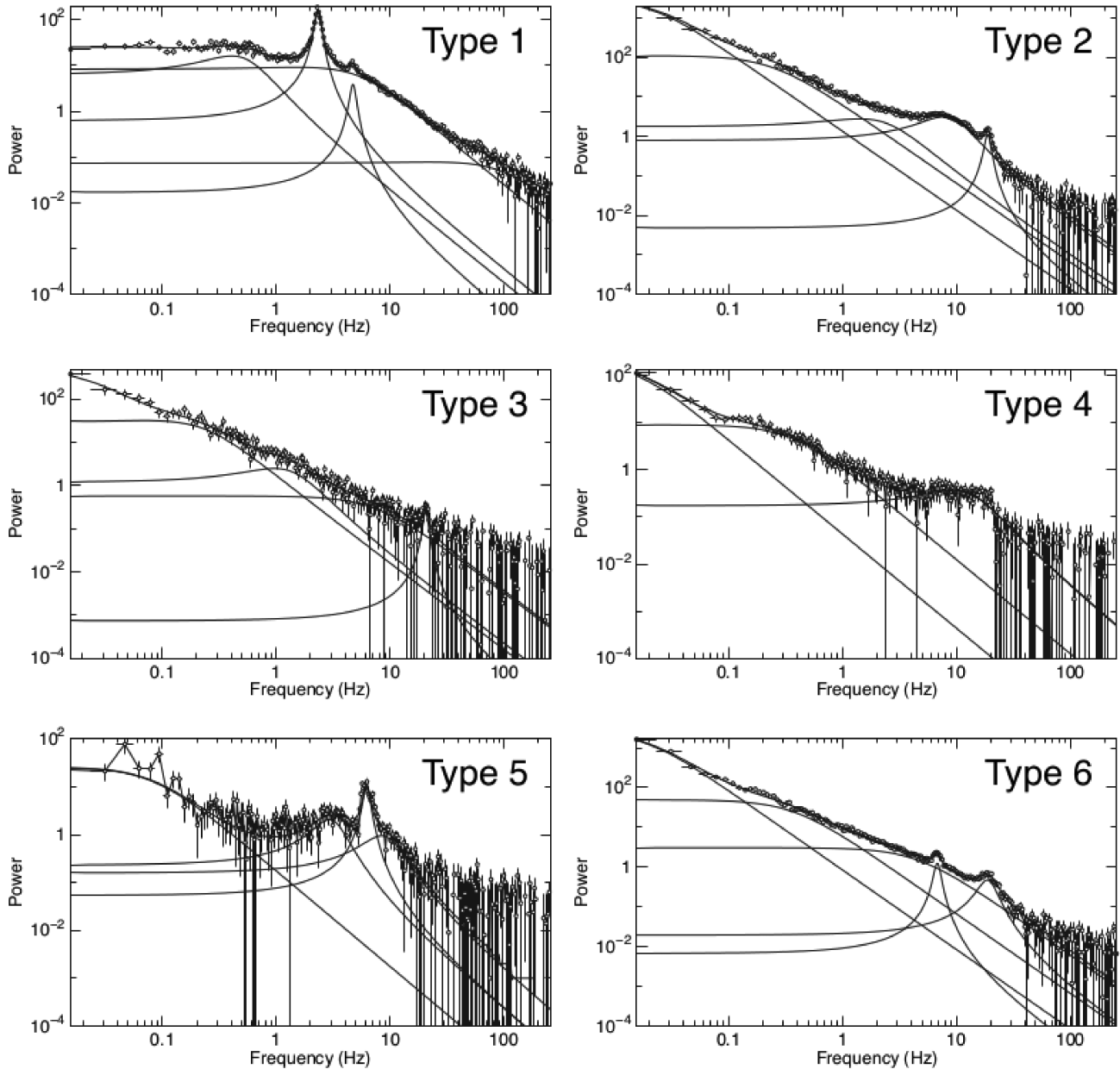


Figure 3. Selection of PDS observed during the 2005 outburst of GRO J1655–40. From PDS types 1–6 we show Obs #27, #36, #55, #65, #28 and #42, respectively. PDS type 6 (Obs #42) shows both a type-B QPO at ~ 6 Hz and a type-C QPO at ~ 18 Hz. The solid curves represent the best fit to each PDS, and its individual components. Notice that the PDS are normalized according to Leahy et al. (1983).

per cent and is narrow ($Q = 6.8$) and moves in frequency on short time-scales (tens of seconds, see Fig. 4). The peak at higher frequency is broader ($Q = 2.4$) and has fractional rms = 1.4 per cent. This peak is weaker than all the other peaks observed at similar frequencies during the 2005 outburst. The spectrogram shows that the width of this peak is intrinsic (i.e. it is not due to fast frequency variations along the observation).⁸ As shown in Fig. 1 (middle and bottom panels), the integrated fractional rms of the whole PDS is significantly lower than the previous and following observations and is close to 5 per cent. The light curve of this observation shows strong flaring on short time-scales. This observation belongs to the

⁸ On a time-scale that is equal or longer than 8 s, i.e. the length of the light-curve stretches used to produce the spectrogram.

ULS and corresponds to the highest count rate observed during this outburst.

3.1.2 Relations between rms and frequency

Following Casella et al. (2005) and Motta et al. (2011), we plot in Fig. 5 the integrated fractional rms of each PDS versus the centroid frequency of the QPO. This is a useful method for differentiating between different types of PDS as the integrated fractional rms is known to correlate well with the frequency of some type of LFQPOs (see e.g. Casella et al. 2005; Kalamkar et al. 2011; Motta et al. 2011; Muñoz-Darias et al. 2011b). In particular, we want to investigate whether the two simultaneous and non-harmonically related peaks detected at the peak of the outburst (type-6 PDS in Fig. 3) can be

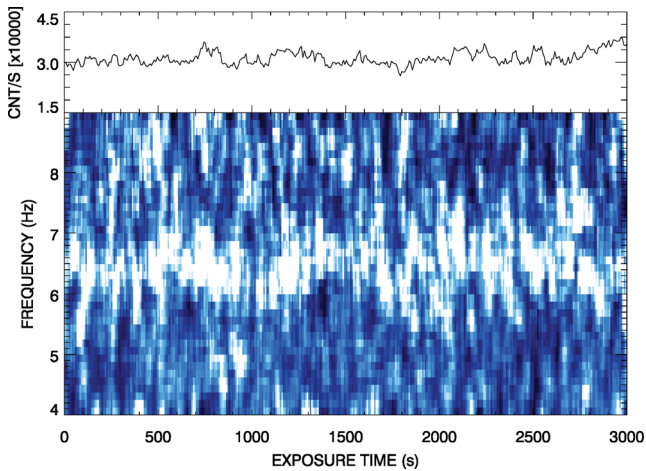


Figure 4. Spectrogram for Obs #42 with a time resolution of 8 s. The Poissonian noise has been subtracted and the power is normalized according to Leahy et al. (1983). Only the first section of the observation is shown.

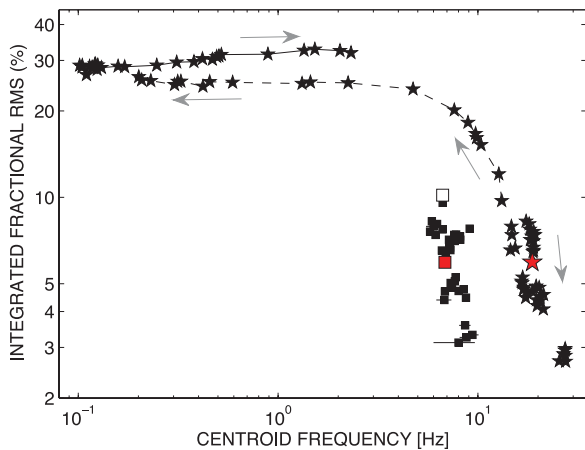


Figure 5. QPO centroid frequency versus integrated (0.1–64 Hz) fractional rms. Each point corresponds to a single *RXTE* observation. The grey arrows mark the QPO frequency evolution. The solid line joins as a function of time the observations showing QPOs in the LHS during the rise of the outburst; the dashed line joins the observations showing QPOs observed in the LHS and the HIMS during the decay of the outburst. We included in the plot both the observations showing a QPO or a peaked noise component. Stars are type-C QPOs and squares type-B QPOs. The red star corresponds to the type-C QPO observed simultaneously with the type-B QPO (red square) in PDS type 6 (Obs #42). The white square corresponds to the type-B QPO observed in PDS type 5 (Obs #28). The black squares correspond to the peaked noise components observed in the ULS (type 2 PDS).

identified as type A, B or C QPOs. For completeness, we also include the points corresponding to peaked noise components observed in PDS types 2 and 4 in Fig. 5.

Fig. 5 shows that two groups of points can be identified and associated with two different types of QPOs. The QPOs observed in the LHS and HIMS during the rise and decay of the outburst are given as function of time, joined by the solid and dashed line, respectively.

(i) A large group of points (black stars) forms a curved track that covers a large frequency (~ 0.1 – 30 Hz) and rms range (3–30 per cent). These points correspond to all the narrow peaks observed in PDS types 1, 2 and 3. The points coming from the rise of the outburst (i.e. along the upper horizontal branch of the HID) are

found at higher integrated fractional rms and at a frequency below ~ 3 Hz. At frequencies above 10 Hz we observe two parallel tracks corresponding to the rise (higher frequency) and decay phase of the ULS.⁹ As Fig. 2 shows, the integrated fractional rms is generally higher during the rise than during the decay of the ULS. According to the ABC classification (Casella et al. 2005), the properties of these peaks, summarized in Table 1, allow us to classify them as type-C QPOs.

(ii) Another group of points (black squares) corresponds to the peaked noise component observed in PDS types 2 and 4. They form a different track along which the frequency weakly anticorrelates with the integrated fractional rms. They are mostly found in the 4–10 per cent integrated fractional rms range, with few points at ~ 3 per cent rms. This track includes the QPO observed in PDS type 5 during the SIMS (Obs #28), marked by a white square in Fig. 5. This last peak can be classified as a type-B QPO according to the ABC scheme.¹⁰ Even though the peaked noise components cannot be properly defined QPOs given their broadness, their properties (see Table 1) suggest similarities to type-B QPOs (see also Section 4).

(iii) The simultaneous QPOs detected at the peak of the outburst at ~ 18 and ~ 7 Hz (type-6 PDS, Obs #42, red star and red square in Fig. 5) fall on the first and second group of points, respectively. The properties of these two peaks allow us to classify them as a type-C QPO and a type-B QPO and show that a type-C and a type-B QPO can be detected simultaneously. The fast frequency variations shown by the peak at lower frequencies (see Fig. 4) further suggest its type-B nature (Nespoli et al. 2003).

4 DISCUSSION

We have analysed 507 *RXTE* observations of the BHXB GRO J1655–40 from its 2005 outburst. We selected all the observations showing LFQPOs in order to investigate their properties and evolution throughout the outburst of the source. This sample consisted of a total of 92 observations.

All LFQPOs could be classified using the ABC scheme and the rms–frequency diagram. This includes those LFQPOs detected in the peculiar PDS types observed in the ULS. Most of the LFQPOs are type-C QPOs. 54 per cent of these are observed during the LHS and the HIMS, showing the same frequency evolution and properties already observed in many other sources (see e.g. GX 339–4, Motta et al. 2011; XTE J1859+226, Casella et al. 2004; MAXI J1659–152, Kalamkar et al. 2011). These QPOs have frequencies between 0.1 and 13 Hz (0.1–3 Hz during the rise phase of the outburst and 0.2–13 Hz during the decay phase). About 30 per cent of the type-C QPOs are found in the ULS at frequencies ranging from ~ 14 to ~ 20 Hz and they usually appear together with a peaked noise component with characteristic frequency of ~ 7 Hz. Interestingly, 16 per cent of the type-C QPOs are observed in the HSS. Type-C QPOs are typically only observed in spectrally harder states, although similar QPOs have also been found in the HSS states of XTE J1550–564 (Homan et al. 2001) and H1743–332 (Homan et al. 2005a). Of the type-C QPOs in the HSS, some are observed very close to the transition from the ULS to the HSS and their frequency is about 20 Hz. Other type-C QPOs are found far from

⁹ We call *rise* of the ULS all the observations of the ULS before Obs #42 (outburst peak) and *decay* all the ULS observations after Obs #48.

¹⁰ Following Casella et al. (2004), the peak takes the form of a type-B *cathedral* in the second section of Obs #28 (see description of type-5 PDS in Section 3.1.1) and of a standard type B in the third section of Obs #28.

Table 1. Summary of the timing properties of GRO J1655–40 during the 2005 outburst. The integrated rms is measured between 0.1 and 64 Hz.

	Type C LHS/HIMS	Type C ULS	Type C HSS	Type B	Peaked component
Frequency (Hz)	0.1–13	14–20	20–28	~7	6–9
Q	3–9	3–11	3–39	3–7	0.4–1.4
Integrated rms (per cent)	10–30	4–8	2–5	5–10	5–10
Noise	Strong flat top	Weak flat top + red	Weak flat top + red	Red	Red

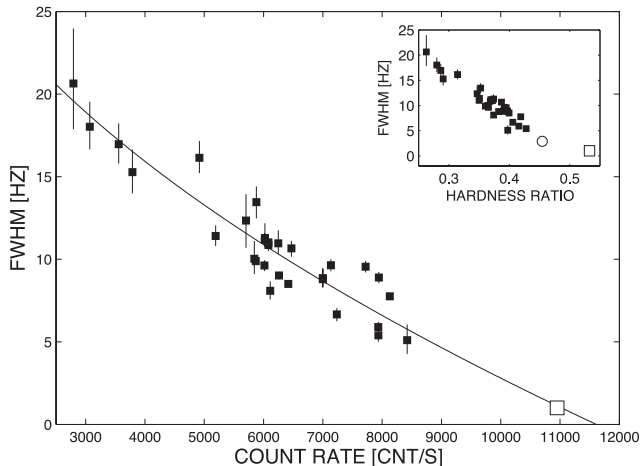


Figure 6. Relation between the FWHM of the peaked noise component in the type 2 PDS and the count rate. The white square represents the type-B QPO observed at the peak of the outburst, simultaneously with the type-C QPO. The relation is well described by a simple polynomial of the form $y = Ax^B + C$, where $A = -0.28$, $B = 0.52$ and $C = 37.29$. The inset shows the FWHM of the noise component as a function of the hardness ratio, where the white dot represents the type-B QPO observed in the SIMS (Obs #28).

the transition and their frequency ranges from 25 to 28 Hz. These QPOs, already reported by Remillard et al. (1999), are among the highest frequency type-C QPOs observed in BHXBs (see Vignarca et al. 2003 for the case of XTE J1748–288). The relations traced out in the frequency–rms diagram by the type-C QPOs observed in the ULS and HSS appear to be continuations of the branch traced out by the HIMS type-C QPOs, although they are shifted towards slightly higher frequencies.

Only one LFQPO displayed all the standard properties of a type-B QPO (i.e. the single SIMS observation during the outburst rise). However, in several observations of the ULS we detected a peaked noise component that shares some of the characteristics of type-B QPOs (although they remain different features, mainly because of their width). This noise component had a characteristic frequency between 6 and 8 Hz and they appear to be related to significant increases in the count rate, similarly to what was observed by e.g. Motta et al. (2011) in the case of type-B QPOs. In the frequency–rms diagram the noise components trace out a track that is well separated from type-C QPOs and that also includes the type-B QPO detected in the SIMS. This track also corresponds to the region where type-B QPOs are usually observed in other sources (see e.g. Casella et al. 2004).

Following the evolution of the peaked noise component along the ULS, it appears that this broad component evolves into a narrow type-B QPO as the count rate rises. Fig. 6 shows the dependence of the peaked noise FWHM on count rate. In the inset of Fig. 6 we show the relation between the FWHM of the peaked noise as

a function of the hardness ratio and we also include the type-B QPO observed at the peak of the outburst (white square). The width of the peaked noise component anticorrelates well with the count rate. As already suggested by Motta et al. (2011), the count rate could rise as a consequence of an increase in the local accretion rate, sudden changes in the geometry or radiative efficiency, or even the appearance of an additional component contributing to the (hard) emission (i.e. jet from the inner regions of the system). Our result might indicate that, at least for what concerns the ULS, if the increase in count rate is strong enough (in other words if the variation in count rate – not the absolute count rate itself – is large enough) the mechanism responsible of the production of type-B QPOs is triggered and the PDS will show either a peaked noise component with characteristic frequency ~6–8 Hz or a proper type-B QPO (as in Obs #28, see below), depending on the count rate level that is reached. At lower count rates the peaked noise component may simply become too broad (and weak) to be detectable. This result shows that under certain conditions sudden count rate increases might constitute a key ingredient in the appearance of type-B QPOs (see also Stiele et al. 2012, for a discussion on the topic).

The timing properties of the ULS, in particular the presence of these type-B-like features in the PDS, suggest that most of the ULS can be regarded as a sort of SIMS covering a higher flux range. We note that in the RID the ULS falls in the region where the SIMS is usually observed (see Belloni et al. 2011; Motta et al. 2011; Muñoz-Darias et al. 2011a). However, Belloni et al. (2011) defines the SIMS on the basis of the presence of a type-B QPO. During the ULS of GRO J1655–40 both type-C QPOs and type-B QPOs (or type-B-like features) appear, as well as a strong red noise. Therefore, even though the ULS shares the characteristics of the SIMS, it still remains significantly different from it.

The unique PDS observed at the peak of the outburst shows two simultaneous non-harmonically related peaks at ~7 and ~18 Hz that we could classify as a type-B and type-C QPO, respectively. The narrow peak at 7 Hz falls on the same relations traced by the peaked noise components in a rms–frequency diagram, while the broader peak at 18 Hz lies on the track formed by the type-C QPOs observed all over the outburst. The peak at 7 Hz also shows significant frequency variations between 6 and 8 Hz on a short time-scale (few seconds) that are typical of type-B QPOs (see Nespoli et al. 2003). It is important to stress that this PDS is observed in the ULS, when the count rate reached its maximum during the 2005 outburst. Reaching a very high count rate (larger than ~10 000 counts s^{-1} PCU $^{-1}$ in the 3–20 keV energy band) and the source being in the ULS could be necessary conditions for the peak at ~7 Hz to become a narrow feature ($Q \leq 2$) and therefore for the appearance of two simultaneous different types of QPOs.

A few additional remarks on the nature of the two simultaneous type-B and type-C QPOs observed at the outburst peak need to be made.

Table 2. Power spectral classification and variability parameters. Only observations with evidence of LFOQPOs or a peaked noise component are listed. The PDS are classified according to the criteria detailed in the text (see Section 3.1.1).

#	Time (MJD)	Obs ID	Hardness ratio	Count rate (counts s ⁻¹)	rms (per cent)	State	PDS type	Type-C QPOs			Type-B QPOs			Peaked noise	
								Frequency (Hz)	FWHM (Hz)	Width (Hz)	Frequency (Hz)	FWHM (Hz)	Frequency (Hz)	FWHM (Hz)	
1	53427	90058-16-05-00	0.837 ± 0.007	48.3 ± 0.2	28.8 ± 0.5	LHS	1	0.105 ^{+0.002} _{-0.003}	0.025 ^{+0.009} _{-0.006}	-	-	-	-	-	-
2	53427.2	90428-01-01-01	0.826 ± 0.006	50.1 ± 0.2	28.4 ± 0.4	LHS	1	0.104 ^{+0.004} _{-0.002}	0.032 ^{+0.006} _{-0.007}	-	-	-	-	-	-
3	53427.9	90058-16-07-00	0.837 ± 0.006	54.2 ± 0.2	29.1 ± 0.3	LHS	1	0.122 ^{+0.002} _{-0.005}	0.035 ^{+0.008} _{-0.007}	-	-	-	-	-	-
4	53428.1	90428-01-01-03	0.830 ± 0.006	53.9 ± 0.2	28.4 ± 0.4	LHS	1	0.117 ^{+0.000} _{-0.003}	<0.004	-	-	-	-	-	-
5	53428.9	90428-01-01-04	0.859 ± 0.004	52.5 ± 0.1	29.1 ± 0.2	LHS	1	0.121 ^{+0.004} _{-0.001}	0.015 ^{+0.004} _{-0.002}	-	-	-	-	-	-
6	53429.7	90428-01-01-02	0.836 ± 0.004	55.7 ± 0.1	28.9 ± 0.2	LHS	1	0.125 ^{+0.002} _{-0.002}	0.019 ^{+0.005} _{-0.002}	-	-	-	-	-	-
7	53431	90428-01-01-05	0.831 ± 0.005	50.8 ± 0.1	28.8 ± 0.3	LHS	1	0.102 ^{+0.009} _{-0.005}	<0.002	-	-	-	-	-	-
8	53431.2	90058-16-06-00	0.834 ± 0.008	51.1 ± 0.2	26.7 ± 0.6	LHS	1	0.110 ^{+0.003} _{-0.003}	0.028 ^{+0.009} _{-0.008}	-	-	-	-	-	-
9	53431.6	90428-01-01-06	0.828 ± 0.008	51.2 ± 0.2	28.3 ± 0.6	LHS	1	0.119 ^{+0.002} _{-0.002}	0.021 ^{+0.006} _{-0.005}	-	-	-	-	-	-
10	53431.7	90428-01-01-07	0.866 ± 0.007	47.8 ± 0.2	28.6 ± 0.4	LHS	1	0.114 ^{+0.003} _{-0.003}	0.042 ^{+0.009} _{-0.007}	-	-	-	-	-	-
11	53431.8	90428-01-01-08	0.869 ± 0.008	49.3 ± 0.2	27.9 ± 0.6	LHS	1	0.124 ^{+0.006} _{-0.007}	0.023 ^{+0.012} _{-0.005}	-	-	-	-	-	-
12	53431.9	90428-01-01-09	0.853 ± 0.004	51.6 ± 0.1	28.3 ± 0.2	LHS	1	0.131 ^{+0.001} _{-0.003}	0.025 ^{+0.004} _{-0.004}	-	-	-	-	-	-
13	53432.8	90428-01-01-10	0.834 ± 0.004	62.2 ± 0.1	28.6 ± 0.2	LHS	1	0.158 ^{+0.002} _{-0.002}	0.030 ^{+0.005} _{-0.004}	-	-	-	-	-	-
14	53433	91404-01-01-00	0.831 ± 0.005	65.2 ± 0.2	28.5 ± 0.3	LHS	1	0.171 ^{+0.003} _{-0.003}	0.018 ^{+0.006} _{-0.003}	-	-	-	-	-	-
15	53433.9	91404-01-01-02	0.822 ± 0.004	87.1 ± 0.2	28.9 ± 0.2	LHS	1	0.247 ^{+0.002} _{-0.003}	0.047 ^{+0.010} _{-0.002}	-	-	-	-	-	-
16	53434.7	91404-01-01-03	0.857 ± 0.005	98.8 ± 0.2	29.6 ± 0.3	LHS	1	0.311 ^{+0.003} _{-0.005}	0.075 ^{+0.008} _{-0.012}	-	-	-	-	-	-
17	53435.6	91404-01-01-01	0.807 ± 0.005	135.2 ± 0.3	29.7 ± 0.3	LHS	1	0.379 ^{+0.003} _{-0.006}	0.08 ^{+0.01} _{-0.01}	-	-	-	-	-	-
18	53436.2	91404-01-01-04	0.804 ± 0.005	157.2 ± 0.4	30.4 ± 0.3	LHS	1	0.418 ^{+0.004} _{-0.010}	0.072 ^{+0.010} _{-0.012}	-	-	-	-	-	-
19	53436.4	91404-01-01-05	0.807 ± 0.006	172.7 ± 0.5	30.3 ± 0.4	LHS	1	0.469 ^{+0.009} _{-0.006}	0.054 ^{+0.008} _{-0.015}	-	-	-	-	-	-
20	53436.7	91702-01-01-00	0.800 ± 0.003	194.1 ± 0.3	30.9 ± 0.1	LHS	1	0.493 ^{+0.005} _{-0.002}	0.107 ^{+0.008} _{-0.007}	-	-	-	-	-	-
21	53437.1	91702-01-01-01	0.809 ± 0.004	207.8 ± 0.4	31.1 ± 0.3	LHS	1	0.506 ^{+0.006} _{-0.007}	0.12 ^{+0.01} _{-0.01}	-	-	-	-	-	-
22	53437.1	91702-01-01-02	0.806 ± 0.005	209.0 ± 0.5	31.4 ± 0.3	LHS	1	0.520 ^{+0.005} _{-0.004}	0.070 ^{+0.018} _{-0.007}	-	-	-	-	-	-
23	53438.1	91702-01-01-03	0.764 ± 0.003	315.7 ± 0.6	31.5 ± 0.2	LHS	1	0.890 ^{+0.004} _{-0.006}	0.16 ^{+0.01} _{-0.01}	-	-	-	-	-	-
24	53438.8	91702-01-01-04	0.748 ± 0.003	415.5 ± 0.6	32.5 ± 0.2	LHS	1	1.351 ^{+0.003} _{-0.006}	0.208 ^{+0.011} _{-0.013}	-	-	-	-	-	-
25	53439.1	91702-01-01-05	0.741 ± 0.003	458.1 ± 0.8	32.8 ± 0.3	LHS	1	1.529 ^{+0.005} _{-0.006}	0.23 ^{+0.01} _{-0.01}	-	-	-	-	-	-
26	53439.6	90704-04-01-01	0.716 ± 0.003	523.6 ± 0.8	32.4 ± 0.2	LHS	1	2.043 ^{+0.004} _{-0.005}	0.27 ^{+0.01} _{-0.01}	-	-	-	-	-	-
27	53439.7	90704-04-01-00	0.704 ± 0.002	561.2 ± 0.8	31.9 ± 0.2	LHS	1	2.317 ^{+0.004} _{-0.005}	0.29 ^{+0.01} _{-0.01}	-	-	-	-	-	-
28	53440.7	91702-01-02-00G	0.455 ± 0.001	2026.0 ± 3.0	9.6 ± 0.0	SIMS	5	-	6.66 ^{+0.03} _{-0.03}	2.9 ^{+0.1} _{-0.1}	-	-	-	-	-
29	53500.8	91702-01-52-03	0.387 ± 0.001	6471.0 ± 10.0	7.1 ± 0.1	ULS	2	18.4 ^{+0.3} _{-0.3}	3.2 ^{+0.6} _{-0.5}	-	-	-	8.2 ^{+0.2} _{-0.2}	10.7 ^{+0.5} _{-0.5}	-
30	53501.6	91702-01-53-01	0.374 ± 0.001	6024.0 ± 9.5	7.8 ± 0.1	ULS	4	-	-	-	-	-	9.1 ^{+0.3} _{-0.3}	13.4 ^{+0.6} _{-0.6}	-

Table 2 – continued

#	Time (MJD)	Obs ID	Hardness ratio	Count rate (counts s ⁻¹)	rms (per cent)	State	PDS type	Type-C QPOs			Type-B QPOs			Peaked noise	
								Frequency (Hz)	FWHM (Hz)	FWHM (Hz)	Frequency (Hz)	Width (Hz)	Frequency (Hz)	FWHM (Hz)	
31	53502.3	91702-01-54-00	0.391 ± 0.001	6261.0 ± 9.6	7.4 ± 0.0	ULS	2	18.9 ^{+0.1} _{-0.1}	2.5 ^{+0.2} _{-0.2}	–	–	7.8 ^{+0.1} _{-0.1}	9.0 ^{+0.2} _{-0.2}		
32	53502.4	91702-01-54-01	0.392 ± 0.001	7136.0 ± 11.0	7.1 ± 0.1	ULS	2	19.0 ^{+0.1} _{-0.1}	3.0 ^{+0.4} _{-0.3}	–	–	7.7 ^{+0.2} _{-0.2}	9.6 ^{+0.4} _{-0.4}		
33	53502.5	91702-01-54-02	0.395 ± 0.001	7721.0 ± 11.8	7.4 ± 0.1	ULS	2	19.3 ^{+0.1} _{-0.1}	2.4 ^{+0.3} _{-0.3}	–	–	7.6 ^{+0.2} _{-0.2}	9.5 ^{+0.3} _{-0.3}		
34	53502.6	91702-01-54-03	0.365 ± 0.001	5844.0 ± 9.2	7.9 ± 0.1	ULS	2	14.7 ^{+0.5} _{-0.5}	9.6 ^{+0.8} _{-0.8}	–	–	6.0 ^{+0.4} _{-0.5}	10.0 ^{+0.9} _{-0.9}		
35	53503.2	91702-01-55-00	0.350 ± 0.001	6082.0 ± 9.6	7.4 ± 0.0	ULS	2	14.8 ^{+0.2} _{-0.2}	8.2 ^{+0.6} _{-0.6}	–	–	6.1 ^{+0.2} _{-0.2}	11.0 ^{+0.5} _{-0.5}		
36	53506.3	91702-01-56-00G	0.416 ± 0.001	7934.0 ± 11.9	8.1 ± 0.0	ULS	2	18.06 ^{+0.02} _{-0.02}	2.2 ^{+0.1} _{-0.1}	–	–	6.2 ^{+0.1} _{-0.1}	5.9 ^{+0.3} _{-0.3}		
37	53507	91702-01-58-03	0.397 ± 0.001	8423.0 ± 13.1	6.7 ± 0.1	ULS	2	19.0 ^{+0.2} _{-0.2}	4.5 ^{+0.5} _{-0.5}	–	–	7.3 ^{+0.2} _{-0.2}	5.1 ^{+0.9} _{-0.8}		
38	53507	91702-01-58-04	0.397 ± 0.001	7945.0 ± 12.3	7.3 ± 0.1	ULS	2	18.6 ^{+0.1} _{-0.1}	2.9 ^{+0.3} _{-0.3}	–	–	8.1 ^{+0.2} _{-0.2}	8.9 ^{+0.3} _{-0.3}		
39	53507.1	91702-01-58-02	0.382 ± 0.001	7001.0 ± 10.9	6.5 ± 0.1	ULS	2	18.9 ^{+0.2} _{-0.2}	3.8 ^{+0.5} _{-0.4}	–	–	6.9 ^{+0.3} _{-0.2}	8.8 ^{+0.7} _{-0.4}		
40	53507.2	91702-01-57-00G	0.406 ± 0.001	7235.0 ± 10.9	7.6 ± 0.0	ULS	2	19.03 ^{+0.07} _{-0.08}	2.3 ^{+0.2} _{-0.2}	–	–	5.8 ^{+0.1} _{-0.2}	6.7 ^{+0.4} _{-0.4}		
41	53507.7	91702-01-59-00	0.419 ± 0.001	8130.0 ± 12.4	7.7 ± 0.1	ULS	2	18.19 ^{+0.05} _{-0.06}	1.9 ^{+0.1} _{-0.1}	–	–	6.6 ^{+0.2} _{-0.2}	7.8 ^{+0.2} _{-0.2}		
42	53508.5	91702-01-58-00	0.533 ± 0.002	10950.0 ± 15.2	4.7 ± 0.0	ULS	6	18.7 ^{+0.1} _{-0.1}	7.8 ^{+0.7} _{-0.5}	6.84 ^{+0.03} _{-0.03}	1.0 ^{+0.1} _{-0.1}	–	–		
43	53509.2	91702-01-59-02	0.428 ± 0.001	7937.0 ± 11.8	8.3 ± 0.0	ULS	2	17.44 ^{+0.02} _{-0.02}	1.62 ^{+0.05} _{-0.05}	–	–	5.9 ^{+0.1} _{-0.1}	5.4 ^{+0.4} _{-0.4}		
44	53509.6	91702-01-58-01	0.400 ± 0.001	6417.0 ± 9.8	7.1 ± 0.0	ULS	2	18.39 ^{+0.05} _{-0.05}	2.2 ^{+0.1} _{-0.1}	–	–	7.2 ^{+0.1} _{-0.1}	8.5 ^{+0.2} _{-0.2}		
45	53510	91702-01-60-02	0.390 ± 0.001	7001.0 ± 10.9	6.5 ± 0.1	ULS	2	19.0 ^{+0.2} _{-0.2}	3.4 ^{+0.5} _{-0.5}	–	–	6.6 ^{+0.2} _{-0.2}	8.9 ^{+0.6} _{-0.6}		
46	53510.1	91702-01-60-00	0.361 ± 0.001	5868.0 ± 9.1	6.7 ± 0.0	ULS	2	15.4 ^{+0.1} _{-0.1}	7.9 ^{+0.3} _{-0.2}	–	–	7.2 ^{+0.1} _{-0.1}	9.9 ^{+0.2} _{-0.2}		
47	53510.3	91702-01-60-01	0.374 ± 0.001	6113.0 ± 9.5	6.6 ± 0.1	ULS	2	14.6 ^{+0.2} _{-0.2}	8.5 ^{+0.5} _{-0.4}	–	–	7.3 ^{+0.2} _{-0.2}	8.1 ^{+0.6} _{-0.5}		
48	53515.2	91702-01-63-01	0.370 ± 0.001	6021.0 ± 9.3	5.1 ± 0.0	ULS	2	16.54 ^{+0.09} _{-0.09}	5.1 ^{+0.3} _{-0.3}	–	–	7.7 ^{+0.1} _{-0.1}	11.2 ^{+0.4} _{-0.4}		
49	53515.7	91702-01-63-00	0.366 ± 0.001	6015.0 ± 9.3	5.3 ± 0.0	ULS	2	16.77 ^{+0.08} _{-0.09}	5.6 ^{+0.2} _{-0.2}	–	–	7.7 ^{+0.1} _{-0.1}	9.6 ^{+0.3} _{-0.3}		
50	53516.2	91702-01-64-02	0.353 ± 0.001	5878.0 ± 9.3	4.5 ± 0.0	ULS	2	17.4 ^{+0.2} _{-0.2}	3.1 ^{+0.6} _{-0.5}	–	–	8.7 ^{+0.3} _{-0.3}	13.5 ^{+0.9} _{-1.0}		
51	53516.4	91702-01-64-03	0.351 ± 0.001	5192.0 ± 8.2	4.8 ± 0.0	ULS	2	17.0 ^{+0.1} _{-0.2}	4.2 ^{+0.4} _{-0.4}	–	–	8.5 ^{+0.2} _{-0.2}	11.4 ^{+0.6} _{-0.6}		
52	53516.6	91702-01-64-00	0.369 ± 0.001	6082.0 ± 9.4	4.9 ± 0.0	ULS	2	17.03 ^{+0.07} _{-0.07}	4.4 ^{+0.2} _{-0.2}	–	–	7.4 ^{+0.1} _{-0.1}	10.9 ^{+0.3} _{-0.3}		
53	53517	91702-01-65-00	0.372 ± 0.001	6249.0 ± 9.7	5.0 ± 0.0	ULS	2	16.7 ^{+0.2} _{-0.2}	5.5 ^{+0.5} _{-0.6}	–	–	7.3 ^{+0.2} _{-0.2}	11.0 ^{+0.8} _{-0.6}		
54	53517.1	91702-01-65-03	0.347 ± 0.001	5704.0 ± 9.2	4.7 ± 0.1	ULS	2	17.5 ^{+0.3} _{-0.3}	4.3 ^{+1.1} _{-0.9}	–	–	8.0 ^{+0.5} _{-0.5}	12 ⁺² ₋₂		
55	53517.7	91702-01-65-01	0.299 ± 0.001	4799.0 ± 8.1	4.5 ± 0.1	HSS	3	20.5 ^{+0.3} _{-0.4}	2.0 ^{+1.0} _{-0.7}	–	–	–	–		
56	53518.1	91702-01-66-03	0.298 ± 0.001	4702.0 ± 7.8	4.9 ± 0.0	HSS	3	20.3 ^{+0.5} _{-0.6}	7 ⁺² ₋₂	–	–	–	–		
57	53518.1	91702-01-66-05	0.298 ± 0.001	4813.0 ± 8.1	4.1 ± 0.1	HSS	3	21.3 ^{+0.4} _{-0.5}	5 ⁺² ₋₁	–	–	–	–		
58	53518.2	91702-01-65-02	0.297 ± 0.001	4779.0 ± 8.0	4.4 ± 0.1	HSS	3	20.4 ^{+0.6} _{-0.6}	7 ⁺² ₋₁	–	–	–	–		
59	53518.7	91702-01-66-01	0.294 ± 0.001	4615.0 ± 7.7	4.6 ± 0.0	HSS	3	21.3 ^{+0.3} _{-0.4}	6 ⁺¹ ₋₁	–	–	–	–		
60	53518.9	91702-01-67-01	0.311 ± 0.001	4725.0 ± 7.7	4.9 ± 0.0	HSS	3	19.6 ^{+0.3} _{-0.3}	3.2 ^{+1.2} _{-0.9}	–	–	–	–		
61	53519	91702-01-67-02	0.315 ± 0.001	4919.0 ± 8.0	4.4 ± 0.0	HSS	2	19.6 ^{+0.5} _{-0.6}	4.6 ^{+1.0} _{-0.8}	–	–	6.8 ^{+0.6} _{-0.6}	16.1 ^{+1.0} _{-0.9}		

Table 2 – *continued*

#	Time (MJD)	Obs ID	Hardness ratio	Count rate (counts s ⁻¹)	rms (per cent)	State	PDS type	Type-C QPOs			Type-B QPOs			Peaked noise	
								Frequency (Hz)	FWHM (Hz)	FWHM (Hz)	Frequency (Hz)	Width (Hz)	Frequency (Hz)	FWHM (Hz)	
62	53519.1	91702-01-67-03	0.296 ± 0.001	4555.0 ± 7.5	4.2 ± 0.0	HSS	3	19.8 ^{+0.6} _{-0.6}	6 ⁺³ ₋₂	–	–	–	–	–	–
63	53526	91702-01-72-00	0.291 ± 0.001	3789.0 ± 6.3	3.6 ± 0.1	HSS	4	–	–	–	–	–	–	8.6 ^{+0.6} _{-0.6}	15 ⁺¹ ₋₁
64	53528.6	91702-01-73-02	0.287 ± 0.001	3557.0 ± 6.0	3.3 ± 0.1	HSS	4	–	–	–	–	–	–	8.8 ^{+0.7} _{-0.7}	17 ⁺¹ ₋₁
65	53531.5	91702-01-77-00	0.280 ± 0.001	3069.0 ± 5.2	3.3 ± 0.1	HSS	4	–	–	–	–	–	–	9.4 ^{+0.7} _{-0.7}	18 ⁺² ₋₁
66	53532.2	91702-01-77-01	0.263 ± 0.001	2794.0 ± 4.9	3.1 ± 0.1	HSS	4	–	–	–	–	–	–	8 ⁺² ₋₂	21 ⁺³ ₋₃
67	53574.4	91702-01-16-10	0.162 ± 0.001	1821.0 ± 3.4	3.0 ± 0.0	HSS	3	27.4 ^{+0.2} _{-0.2}	1.8 ^{+0.6} _{-0.5}	–	–	–	–	–	–
68	53575.6	91702-01-17-10	0.172 ± 0.001	1868.0 ± 3.4	2.7 ± 0.1	HSS	3	27.51 ^{+0.13} _{-0.10}	0.7 ^{+0.3} _{-0.2}	–	–	–	–	–	–
69	53580.4	91702-01-21-10	0.161 ± 0.001	1823.0 ± 3.4	2.8 ± 0.1	HSS	3	27.4 ^{+0.4} _{-0.4}	3.1 ^{+0.9} _{-0.7}	–	–	–	–	–	–
70	53583.4	91702-01-24-10	0.158 ± 0.001	1716.0 ± 3.2	2.9 ± 0.1	HSS	3	27.3 ^{+0.1} _{-0.1}	1.0 ^{+0.4} _{-0.4}	–	–	–	–	–	–
71	53585.4	91702-01-25-11	0.157 ± 0.001	1678.0 ± 3.2	2.9 ± 0.1	HSS	3	26.9 ^{+0.2} _{-0.2}	1.7 ^{+0.5} _{-0.4}	–	–	–	–	–	–
72	53593.3	91702-01-32-10	0.147 ± 0.001	1514.0 ± 2.9	2.7 ± 0.1	HSS	3	25.6 ^{+0.4} _{-0.2}	1.4 ^{+0.9} _{-0.4}	–	–	–	–	–	–
73	53628.2	91702-01-76-00	0.367 ± 0.001	341.5 ± 0.6	9.7 ± 0.1	HIMS	1	13.17 ^{+0.06} _{-0.06}	1.5 ^{+0.2} _{-0.2}	–	–	–	–	–	–
74	53628.6	91702-01-76-01	0.406 ± 0.002	332.2 ± 0.8	12.1 ± 0.4	HIMS	1	12.74 ^{+0.10} _{-0.10}	1.7 ^{+0.5} _{-0.4}	–	–	–	–	–	–
75	53628.9	91702-01-71-03	0.500 ± 0.003	316.1 ± 0.7	16.1 ± 0.3	HIMS	1	9.86 ^{+0.02} _{-0.02}	5 ⁺² ₋₂	–	–	–	–	–	–
76	53629	91702-01-71-04	0.487 ± 0.002	321.6 ± 0.7	15.3 ± 0.3	LHS	1	10.37 ^{+0.02} _{-0.02}	0.38 ^{+0.07} _{-0.06}	–	–	–	–	–	–
77	53629.4	91702-01-79-01	0.514 ± 0.002	310.8 ± 0.6	16.7 ± 0.2	LHS	1	9.74 ^{+0.01} _{-0.02}	0.56 ^{+0.04} _{-0.04}	–	–	–	–	–	–
78	53630.5	91702-01-79-00	0.543 ± 0.002	275.0 ± 0.4	18.2 ± 0.1	LHS	1	8.93 ^{+0.01} _{-0.01}	0.90 ^{+0.03} _{-0.03}	–	–	–	–	–	–
79	53631.5	91702-01-80-00	0.584 ± 0.002	237.7 ± 0.4	20.1 ± 0.1	LHS	1	7.62 ^{+0.01} _{-0.01}	0.89 ^{+0.03} _{-0.03}	–	–	–	–	–	–
80	53632.5	91702-01-80-01	0.666 ± 0.002	194.8 ± 0.3	23.8 ± 0.1	LHS	1	4.742 ^{+0.008} _{-0.009}	0.62 ^{+0.03} _{-0.03}	–	–	–	–	–	–
81	53633.5	91702-01-81-00	0.735 ± 0.004	148.4 ± 0.3	25.0 ± 0.1	LHS	1	2.24 ^{+0.01} _{-0.02}	0.58 ^{+0.05} _{-0.05}	–	–	–	–	–	–
82	53634.1	91702-01-80-02	0.749 ± 0.006	126.2 ± 0.4	25.2 ± 0.3	LHS	1	1.45 ^{+0.02} _{-0.02}	0.30 ^{+0.05} _{-0.05}	–	–	–	–	–	–
83	53634.3	91702-01-81-01	0.757 ± 0.004	119.3 ± 0.2	24.9 ± 0.2	LHS	1	1.318 ^{+0.006} _{-0.011}	0.23 ^{+0.02} _{-0.03}	–	–	–	–	–	–
84	53635.5	91702-01-81-02	0.781 ± 0.004	84.7 ± 0.2	25.2 ± 0.2	LHS	1	0.592 ^{+0.004} _{-0.005}	0.06 ^{+0.01} _{-0.01}	–	–	–	–	–	–
85	53636.2	91702-01-87-03	0.786 ± 0.005	71.1 ± 0.2	25.3 ± 0.4	LHS	1	0.45 ^{+0.01} _{-0.01}	0.07 ^{+0.04} _{-0.01}	–	–	–	–	–	–
86	53636.5	91702-01-82-00	0.865 ± 0.005	58.1 ± 0.2	24.4 ± 0.5	LHS	1	0.421 ^{+0.009} _{-0.009}	0.10 ^{+0.04} _{-0.01}	–	–	–	–	–	–
87	53637.2	91704-01-01-00	0.787 ± 0.006	58.9 ± 0.2	25.4 ± 0.4	LHS	1	0.327 ^{+0.008} _{-0.004}	0.06 ^{+0.02} _{-0.01}	–	–	–	–	–	–
88	53637.2	91704-01-01-01	0.791 ± 0.004	57.5 ± 0.1	25.3 ± 0.2	LHS	1	0.315 ^{+0.003} _{-0.002}	0.051 ^{+0.006} _{-0.004}	–	–	–	–	–	–
89	53637.5	91704-01-01-02	0.796 ± 0.006	56.0 ± 0.2	24.7 ± 0.4	LHS	1	0.303 ^{+0.003} _{-0.008}	0.043 ^{+0.024} _{-0.009}	–	–	–	–	–	–
90	53638.4	91702-01-86-00	0.788 ± 0.005	47.7 ± 0.1	25.5 ± 0.4	LHS	1	0.231 ^{+0.003} _{-0.007}	0.063 ^{+0.029} _{-0.011}	–	–	–	–	–	–
91	53639.1	91702-01-86-01	0.806 ± 0.007	42.2 ± 0.2	25.7 ± 0.5	LHS	1	0.208 ^{+0.007} _{-0.003}	0.045 ^{+0.027} _{-0.005}	–	–	–	–	–	–
92	53639.2	91702-01-86-04	0.821 ± 0.008	41.4 ± 0.2	26.3 ± 0.7	LHS	1	0.201 ^{+0.004} _{-0.007}	0.062 ^{+0.023} _{-0.017}	–	–	–	–	–	–

(i) The low-frequency noise associated with the two simultaneous peaks is of a hybrid kind: power-law noise plus a weak flat-topped noise superimposed on it. The red noise is typically associated with type-A and type-B QPOs, while flat-topped noise is usually observed together with type-C QPOs.

(ii) The QPO classified as type C is broader than the other type-C QPOs observed during the outburst (see Table 2, e.g. Obs #41 and #42) and it appears similar to a type-A QPO in shape (see Casella et al. 2004, 2005; Motta et al. 2011). Nevertheless, its properties indicate a type-C nature. In addition, we note that type-C QPOs sometimes become broad and quite weak close to the transition between HIMS and SIMS (see Homan et al., in preparation).

(iii) The integrated fractional rms of this PDS is ~ 5 per cent, which is significantly lower than the usual values seen for type-C QPOs (i.e. > 5 per cent), but consistent with the typical rms level at which type-A QPOs are observed (Motta et al. 2011). This observation is very close to the rms limit associated with the transition from SIMS to HSS (~ 5 per cent rms; see Belloni et al. 2011; Muñoz-Darias et al. 2011a).

Type-A and -C QPOs usually have significantly different properties, in particular for what concerns the broad-band noise that comes with these two types of oscillations; (strong) flat-top noise in the case of type-C QPOs and weak power law or weak peaked noise in the case of type-A QPOs. Motta et al. (2011) showed that type-C and -A QPOs share several properties, the most important one being the fact that their frequencies follow the same relations as a function of the flux of the hard component in the energy spectrum. Therefore, they could share a common origin. In this scenario type-A and -C QPOs would be two different effects of a same physical process. Some of the QPOs in GRO J1655–40 (e.g. the peak at ~ 18 Hz at the outburst peak and a few of the type-C QPOs observed in HSS) show characteristics of both type-C and -A QPOs. As discussed in Section 3.1.1, the QPO at 18 Hz observed simultaneously with the type-B QPO at the outburst peak shows the behaviour of a type-C QPO (e.g. it follows the typical type-C frequency evolution), but the characteristics detailed above (e.g. intrinsic broadness and faintness) suggest that it also shows some of the properties of type-A QPOs. Therefore it might be regarded as a link between type-C and type-A QPOs. In the PDS of GRO J1655–40, we did not detect any type-A QPOs (usually observed in the SIMS), but we found several type-C QPOs associated with an ‘hybrid’ noise (formed by a weak flat-top noise component superimposed to a strong red noise) in HSS, where it is uncommon to observe any type of QPO. Unfortunately the complete lack of type-A QPOs in the PDS of GRO J1655–40 makes impossible a detailed comparison between type-A QPOs and HSS type-C QPOs.

The simultaneous detection of a type-B and a type-C QPO proves that at least two intrinsically different types of LFQPOs exist. The fact that two different types of QPOs can be observed at the same time is of great importance to understand the origin of LFQPOs, since it indicates that they are likely produced by different physical processes which, under opportune conditions, can take place simultaneously.

Stella & Vietri (1998) originally proposed that the Lense–Thirring precession could be the mechanism producing LFQPOs in accreting compact objects. Ingram, Done & Fragile (2009) developed this model in the truncated disc framework (see Done et al. 2007), where the precession of a hot flow in the inner regions of the accretion plasma is the origin of type-C QPOs. Motta et al. (2011) suggested that the same mechanism could also explain the presence of type-A QPOs in the emission of black hole candidates (BHCs).

The results reported here essentially rule out the possibility that type-B QPOs could arise from the same physical phenomenon – regardless of what this phenomenon is – supporting what was already suggested by Motta et al. (2011).

Our results show that the presence of two simultaneous type-C and -B QPOs in the emission from GRO J1655–40 is possibly related to the ULS. Similar cases have been found in the 1995 outburst of GRO J1655–40 as well as in the PDS of the BHC XTE J1550–564 and H1743–322 (Motta et al., in preparation). In all these cases simultaneous type-B and -C QPOs are observed at the highest count rate end of the ULS. However, only GRO J1655–40 showed a sufficiently long and bright ULS associated with a standard outburst evolution, that offered the possibility to study in detail the evolution of the timing properties. The investigation of the physical circumstances underlying the presence of two simultaneous different types of LFQPOs is beyond the scope of this work. However, a detailed spectral analysis of the ULS in the 2005 outburst of GRO J1655–40 as support of the timing analysis will be presented in a forthcoming work.

5 SUMMARY AND CONCLUSIONS

We report on a timing analysis of 507 *RXTE* observations of the BHC GRO J1655–40 during its 2005 outburst. We selected all the observations in which a LFQPO was observed and we classified the QPOs according to the ABC scheme.

Our sample included a unique observation whose PDS show two simultaneous, non-harmonically related QPOs which we identified as a type-B and a type-C QPO. This result proves that in the emission of GRO J1655–40 at least two of the three known types of LFQPOs are intrinsically different phenomena and likely produced by different physical processes. This could be valid by extension for other sources as well. Our results also show that the requirements for the appearance of a type-B or a type-C QPO are usually mutually exclusive, but under opportune conditions, these two different types of QPO can be observed simultaneously. Of the two simultaneous QPOs, the one classified as a type C shows some resemblance to a type-A QPOs as well. This weakens the distinction between type-C and type-A QPOs, supporting the hypothesis that these two types of LFQPOs may share a similar physical origin.

We also investigated the nature of the broad peaked noise component observed in most PDS of the ULS observations and we found that it is linked to the type-B QPO observed simultaneously with the type-C QPO at the brightest observation of the outburst. This peaked component is present during local maxima in the light curve. Its width correlates well with the count rate, to the extent that it evolves into a narrow type-B QPO when the count rate is higher than a certain value.

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