THE ATACAMA COSMOLOGY TELESCOPE: DETECTION OF SUNYAEV-ZEL'DOVICH DECREMENT IN GROUPS AND CLUSTERS ASSOCIATED WITH LUMINOUS RED GALAXIES

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ABSTRACT

We present a detection of the Sunyaev-Zel'dovich (SZ) decrement associated with the Luminous Red Galaxy (LRG) sample of the Sloan Digital Sky Survey. The SZ data come from 148 GHz maps of the equatorial region made by the Atacama Cosmology Telescope (ACT). The LRG sample is divided by luminosity into four bins, and estimates for the central Sunyaev-Zel'dovich temperature decrement are calculated through a stacking process. We detect and account for a bias of the SZ signal due to weak radio sources. We use numerical simulations to relate the observed decrement to Y_{200} and clustering properties to relate the galaxy luminosity bins to mass. We also use a relation between BCG luminosity and cluster mass based on stacked gravitational lensing measurements to estimate the characteristic halo masses. The masses are found to be in the range $\sim 10^{13} - 10^{14} h^{-1} M_\odot$, a lower range than has been previously probed.

Subject headings: cosmology:cosmic microwave background — cosmology:observations — galaxies:clusters — Sunyaev-Zel'dovich Effect

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1. INTRODUCTION

New high resolution measurements of the sky at millimeter wavelengths are enabling the use of the thermal Sunyaev-Zel'dovich (SZ) effect as a precise cosmological tool. The effect arises from the inverse Compton scattering of the cosmic microwave background (CMB) and high energy electrons, usually occurring in the hot intra-cluster medium (ICM) of galaxy clusters (Zel'dovich & Sunyaev 1969; Sunyaev & Zel'dovich 1970). For frequencies below 218 GHz, the distortion manifests itself as an arcminute-scale temperature decrement along the line of sight to a cluster. The amplitude of this decrement is nearly independent of redshift, making the SZ effect particularly useful for examining the high redshift universe. The effect can also be used to explore the role of baryonic physics in cluster evolution, as its amplitude is proportional to the thermal pressure of the ICM integrated along the line of sight. For a cluster in hydrostatic equilibrium, it is expected that the integrated SZ signal scales with total cluster mass. For a full description of the SZ effect, see the review articles by Rephaeli (1995), Birkinshaw (1999), and Carlstrom et al. (2002).

Analyses of large X-ray and optically selected cluster samples have recently illustrated the cosmological potential of cluster surveys (e.g., Vikhlinin et al. 2009; Mantz et al. 2010; Rozo et al. 2010). The emergence of large-area SZ cluster surveys, such as the South Pole Telescope (SPT; Carlstrom et al. 2009) and the Atacama Cosmology Telescope (ACT; Fowler et al. 2007; Swetz et al. 2010), will supplement this previous work. Furthermore, upcoming results from the *Planck* satellite are expected to include a large number of SZ-detected clusters (Bartlett et al. 2008). The SPT collaboration reported its first SZ detections in Staniszewski et al. (2009) and has since identified twenty-two cluster candidates in Vanderlinde et al. (2010). Twenty-one of these clusters were optically confirmed in High et al. (2010). The X-ray properties of the SPT sample and the derived SZ/X-ray relations are presented in a follow-up study (Andersson et al. 2010). The ACT collaboration presented its first cluster detections in Hincks et al. (2010) and has since reported on twenty-three clusters in Marriage et al. (2010b), all of which have been optically confirmed (Menanteau et al. 2010). The cosmological implications of the ACT sample are discussed in Sehgal et al. (2010b).

In order to fully use cluster surveys for cosmological purposes, independent and robust estimates of the cluster masses are needed. Several studies have explored the relation between the integrated SZ signal Y and total cluster mass M for massive clusters ($M > 10^{14} M_\odot$) (i.e., Benson et al. 2004; Bonamente et al. 2008; Melin et al. 2010; Huang et al. 2010; Plagge et al. 2010). The integrated Y parameter is defined as

$$Y = \int_\Omega y(\hat{n})d\Omega, \qquad (1)$$

where y is the usual Compton y-parameter and Ω is the solid angle of the cluster. Assuming that thermal energy results solely from gravitational collapse, it is possible to derive self-similar scalings between observables (e.g., Y) and cluster mass. Furthermore, observations show that the SZ-mass relation is relatively insensitive to the specifics of cluster physics (i.e., cooling, AGN feedback) (Bonamente et al. 2008; Sehgal et al. 2010b).

In this work, we report on the stacking of a subset of the ACT 2009 equatorial data at the positions of Luminous Red Galaxies (LRGs) measured by the Sloan Digitial Sky Survey (SDSS). Various studies have examined the clustering properties of the SDSS LRGs (i.e., Zheng et al. 2009; Reid & Spergel 2009) and have found that most of these galaxies reside in halos of typical mass $M \sim 10^{13} - 10^{14} h^{-1} M_\odot$. This paper describes the stacking process and the detection of SZ signal after binning the sample by luminosity. Estimates for the characteristic halo mass of each luminosity bin are obtained using two methods: gravitational lensing mass measurements and an empirical model for halo bias as a function of luminosity (Tegmark et al. 2004; Reyes et al. 2008; Zehavi et al. 2005) that describes the LRG sample well, as shown in Percival et al. (2007). Finally, we present an initial analysis of the relation between integrated SZ signal and halo mass. Unless otherwise stated, we assume a flat ΛCDM cosmology with $\Omega_m = 0.264$, $\Omega_\Lambda = 0.736$, $\sigma_8 = 0.80$ and $H_\circ = 100h$ km s^{-1} Mpc^{-1} with $h = 0.71$ (Komatsu et al. 2009).

This paper is organized as follows. In Section 2, we describe the LRG sample and the ACT map. In Section 3, we describe the process of filtering, and in Section 4 we discuss our methods for binning and stacking the ACT data. We present the stacking results in Section 5 and examine the SZ-mass scaling relation in Section 6. We conclude in Section 7.

2. DATA

In this section, we describe the subset of the ACT data used in this analysis (Sec. 2.1) and the LRG sample from the SDSS (2.2).

2.1. *ACT Data*

ACT is a six-meter telescope operating in the Atacama Desert of Chile at an altitude of 5200 meters. The telescope site was chosen for its dry climate as well as the ability to observe in both the northern and southern skies. The telescope has three 1024-element arrays of transition edge sensors, one each operating at 148 GHz, 218 GHz, and 277 GHz. For a more detailed introduction to the ACT instrument, observations, and data reduction, see Fowler et al. (2007), Swetz et al. (2010) and Das et al. (2010). The science release from the ACT 2008 survey of a region south of the Galactic equator includes results on the CMB power spectrum and related parameter constraints (Fowler et al. 2010; Das et al. 2010; Dunkley et al. 2010; Hajian et al. 2010). Results are also presented on compact millimeter sources (Marriage et al. 2010a) and clusters (Marriage et al. 2010b; Menanteau et al. 2010; Sehgal et al. 2010b). The present study uses data at 148 GHz from a subregion of the 2009 equatorial survey. The subregion covers 268 square-degrees and lies between right ascensions $21^h 20^m$ and $03^h 40^m$ and declinations $-01°30'$ and $01°30'$. This region is coincident with the SDSS Stripe 82, which contains a rich multi-wavelength data set. Figure 1 is a map of the sensitivity across the subregion of the study, along with the locations of the LRGs. The map has been match-filtered for
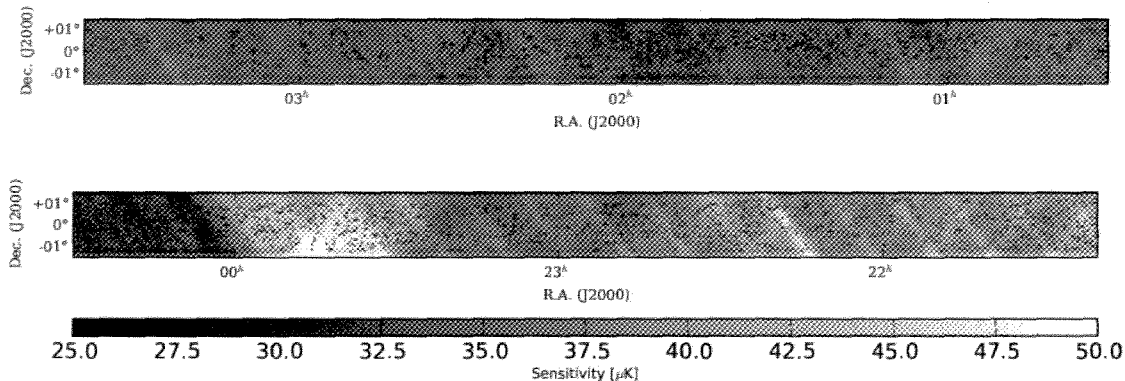
FIG. 1.— Sensitivity map with LRG locations. The map shows the sensitivity over the subset of the ACT 2009 148 GHz equatorial data considered for this study. The gray-scale encodes the noise rms in μK of a map match-filtered for cluster detection. The median noise in the map is 34 μK in CMB temperature units.

cluster detection as described in Section 3. The median rms in the filtered map is 34 μK in CMB temperature units.

2.2. LRG Sample

The LRG sample used in this study is from the SDSS, a survey that has performed five-band ($ugriz$) photometry (Fukugita et al. 1996) using a specialized wide-field camera (Gunn et al. 1998) and multi-object spectroscopy using a pair of spectrographs. The survey has imaged one quarter of the sky at high Galactic latitude and conducted spectroscopic follow up observations of approximately one million of the detected objects (Eisenstein et al. 2001; Richards et al. 2002; Strauss et al. 2002). All data are processed automatically by pipelines that detect and measure photometric properties of sources (Lupton et al. 2001). The SDSS has had seven major data releases (Abazajian et al. 2003, 2004, 2005; Adelman-McCarthy et al. 2006, 2007, 2008; Abazajian et al. 2009).

The LRG catalog used in the current work is drawn from the sample prepared by Kazin et al. (2010), which is taken from the seventh data release (DR7) of the SDSS. The data set was first developed by Eisenstein et al. (2001) and is publicly available.[1] The DR7 LRG sample contains ~110,000 galaxies and extends to redshift $z \simeq 0.5$. The subset used in this study contains 2681 LRGs that fall within the 268 square-degree ACT map. The subset has a spectroscopic redshift range $0.16 < z < 0.47$ and has rest frame g-band absolute magnitudes $-23.2 < M_g < -21.2$. K corrections (Blanton & Roweis 2007) have been applied to the galaxies. The SDSS catalog of LRGs serves as a good tracer of large scale structure, as LRGs probe a larger volume than other tracers since they are more luminous than most galaxies. They also exhibit a distinct 4000 Å break in their spectral energy distributions, making redshift determination easier than for other galaxies.

3. FILTERING THE ACT DATA

To maximize the SZ signal of the clusters associated with the LRGs, we employ a matched fil-

[1] http://cosmo.nyu.edu/~eak306/SDSS-LRG.html

ter (Haehnelt & Tegmark 1996; Herranz et al. 2002; Melin et al. 2006). The filtering scheme used here closely follows Marriage et al. (2010a), with exceptions noted below. At a position x in the map, we model the temperature measured as the sum of the SZ signal plus the other components in the map (CMB fluctuations, atmospheric fluctuations, instrument noise and point sources):

$$\delta T(x) = \sum_i \delta T_{\circ,i} b_i(x - x_i) + \delta T_{\text{other}}(x), \qquad (2)$$

where $\delta T_{\circ,i}$ is the peak amplitude of the ith LRG and b_i is the unit-normalized ACT 148 GHz beam function, taken to be isotropic with a FWHM of 1.4' (Hincks et al. 2010). We first mask the brightest sources in the map, which amount to 146 objects. Only pixels associated with a source with at least SNR > 5 are masked. These are predominantly radio sources (Marriage et al. 2010a). None of these sources lies within 0.5' of an LRG position. Before applying the matched filter, we multiply the map, pixel-wise, by the number of observations per pixel normalized by the maximum number of observations in any single pixel, $\sqrt{N_{\text{obs}}(x)/N_{\text{obs,max}}}$. This weighting accounts for local changes in the amplitude of the white noise and is equivalent to down-weighting pixels with a high white noise rms. This weight scheme is also used during the stacking process when determining the matched filter decrement $\delta T_{\circ}$ (see Section 4.2). Next, the map is filtered in Fourier space using a matched filter:

$$\delta T_{\text{filt}}(k) = \frac{\tilde{b}^{\star}(k) \, |\delta \tilde{T}_{\text{other}}(k)|^{-2} \, \delta \tilde{T}(k)}{\int \tilde{b}^{\star}(k') \, |\delta \tilde{T}_{\text{other}}(k')|^{-2} \, \tilde{b}(k') dk'}. \qquad (3)$$

Since the power spectra of the atmosphere and the CMB scale as l^{-4} for $l > 1000$, the filter is approximately a fourth derivative operator convolved with the beam. The choice of the beam function b for the spatial profile of the cluster differs from the method used in Marriage et al. (2010b), which employs a β-model for the cluster profile. By comparisons with simulations, we expect that the ACT 148 GHz beam function will be a good match for the spatial profile of clusters and hence adopt it in the matched filter. It is particularly effective at reducing

TABLE 1
LUMINOSITY BINS AND HALO MASSES FOR STACKED LRGS

Bin	$N_{\rm bin}$	$\langle L_{0.1r} \rangle$ $10^{10} L_\odot$	$L_{0.1r}$ Range $10^{10} L_\odot$	$\langle z \rangle$	$M_{200\bar{\rho}}$[a] $10^{14} M_\odot$	$M_{200\bar{\rho}}$[b] $10^{14} M_\odot$	$M_{200\bar{\rho}}$[c] $10^{14} M_\odot$
1	26	30.1	25.9 - 44.9	0.39	2.74 ± 0.74	5.20 ± 1.23	2.59 ± 0.87
2	60	23.6	21.8 - 25.3	0.40	1.78 ± 0.13	3.57 ± 0.23	2.05 ± 0.43
3	674	16.7	13.8 - 21.8	0.37	1.00 ± 0.22	2.17 ± 0.41	1.43 ± 0.31
4	1921	10.0	1.1 - 13.8	0.30	0.43 ± 0.14	1.07 ± 0.29	0.82 ± 0.22

[a]Determined from halo bias using $b_\star = 0.99$.
[b]Determined from halo bias using $b_\star = 1.30$.
[c]Determined from lensing measurements in Reyes et al. (2008).

scatter from CMB noise, to which larger filters are more sensitive.

Rather than explicitly modeling the noise contribution $\delta T_{\rm other}(\boldsymbol{x})$, we use the original input map with sources masked for the power spectrum of $\tilde{\delta T}_{\rm other}(\boldsymbol{k})$ in equation 3. As the model for $\delta T_{\rm other}(\boldsymbol{x})$ is not well-conditioned, we smooth $|\tilde{\delta T}_{\rm other}(\boldsymbol{k})|^{-2}$ with a Gaussian filter ($\sigma = 5$ Fourier space pixels). We also set outlying Fourier space pixels with amplitude greater than 10 times the median pixel value to the median value. We also include an additional low-ℓ filter that grows from zero at $\ell = 0$ to unity at $\ell = 1200$ as $\sin^5(\pi\ell/2400)$. This filter accounts for any modification of CMB statistics due to multiplying by the non-uniform weight map and makes our result more robust to any potential non-Gaussianity in the large-scale noise that might arise from residual atmospheric noise in the maps. After finishing the filtering process, we do not include any sources that fall within $10'$ of the map edge in order to mitigate noise from edge effects. This reduces the map size from 284 square degrees to 268 square-degrees. The entirety of the LRG sample discussed above (2681 objects) falls within the bounds of the reduced map, leading to roughly 10 LRGs per square-degree.

4. STACKING METHODS

While ACT can easily detect massive clusters ($M > 10^{15} M_\odot$) (Marriage et al. 2010b), the CMB fluctuations and detector noise limit the direct detection of less massive clusters. Since the mean of both the CMB fluctuations and the atmospheric noise should be zero, we stack the LRG cluster submaps to enable the detection of weaker SZ signals.

4.1. Luminosity Bins

Before stacking the LRGs, we bin the sample by r-band luminosity K corrected to $z = 0.10$, $L_{0.1r}$. Table 1 shows the four luminosity bins used in the analysis. This table also lists mass estimates for each bin, as described in Section 5.2. The luminosities have been computed using the Galactic extinction-corrected (using maps from Schlegel et al. 1998) Petrosian magnitudes from the DR7 CAS and their measured errors. They have been K corrected to $z = 0.10$ using the kcorrect_v4_2 software package (Blanton & Roweis 2007). There have been no added evolutionary corrections. We also refer to the K corrected $^{0.1}r$-band absolute magnitude of an LRG as $M_{0.1r}$. Following Blanton et al. (2003),

$$M_{0.1r} = m_r - DM(z, \Omega_m, \Omega_\Lambda, h) - K_{0.1rr}(z), \qquad (4)$$

where $DM(z, \Omega_m, \Omega_\Lambda, h)$ is the distance modulus in a ($\Omega_m = 0.3, \Omega_\Lambda = 0.7, h = 0.71$) cosmology and $K_{0.1rr}(z)$ is the K correction from the r-band of a galaxy at redshift z to the $^{0.1}r$-band. Luminosities are reported in solar luminosities, where we have used $M_{\odot,0.1r} = 4.76$ (Blanton et al. 2003). These definitions of luminosity and magnitude will facilitate the estimation of halo masses for our binned data in Section 5.2.

4.2. Algorithm

We stack filtered submaps with area $10' \times 10'$ centered on each LRG in a given luminosity bin, and then estimate the matched filter decrement $\delta T_\circ$ from the central temperature value in the stacked map. Recall that the filtered map is weighted by the number of observations per pixel normalized by the maximum number of observations in any single pixel. To minimize the errors induced by the $0.5'$ pixel size of the ACT maps (Marriage et al. 2010a), the pixel size of the submap centered on each LRG is decreased using Fourier interpolation. For the present analysis, the pixel size is decreased to $0.03125'$ and the submaps are recentered. Before determining $\delta T_\circ$, the stacked map for a given bin is convolved with the ACT 148 GHz beam function b to recover decrement information from the surrounding pixels. Lastly, we sum over the five pixels within $4''$ of the centroid of the stacked map and divide by the sum of the weights to compute $\delta T_\circ$. We compute the weight for each submap in the stack before decreasing the pixel size by using the value of $\sqrt{N_{\rm obs}/N_{\rm obs,max}}$ at its central $0.5'$ pixel.

5. STACKING RESULTS

5.1. Matched Filter Decrement $\delta T_\circ$

Tables 2 and 3 contain estimates for the matched filter decrement $\delta T_\circ$ obtained for each of the four luminosity bins. Due to radio contamination (see discussion below), we report results for the full LRG sample as well as a radio-quiet subset, which excludes approximately 10% of the LRG sample. The value for $\delta T_\circ$ reported in Tables 2 and 3 is the temperature decrement in the ACT map, recovered from the matched filter. Figure 2 shows $7' \times 7'$ stacked submaps for each luminosity bin for both LRG samples. Stacked maps obtained when using random LRG positions are also shown, and it is apparent that there is no clear detection in these maps, as expected.

In addition to the central peak in the stacked images, several other significant peaks are evident, particularly the deep cold spot at the upper edge of the stacked map for Bin 2, which has a depth of around 24 μK and is

TABLE 2
CENTRAL SZ TEMPERATURE DECREMENT AND INTEGRATED SZ SIGNAL FOR ALL STACKED LRGS, BINNED BY $L_{0.1r}$

Bin	N_{bin}	$\delta T_\circ$ μK	SNR	$Y_{200\bar\rho}$ arcmin^2	$Y_{200\bar\rho}\ d_A^2\ E(z)^{-2/3}$ Mpc^2
1	26	-27.5 ± 11.3	2.4	$(1.5 \pm 0.5) \times 10^{-4}$	$(1.1 \pm 0.4) \times 10^{-5}$
2	60	-9.9 ± 5.3	1.9	$(4.7 \pm 2.1) \times 10^{-5}$	$(3.4 \pm 1.8) \times 10^{-6}$
3	674	-4.7 ± 1.5	3.1	$(1.8 \pm 0.5) \times 10^{-5}$	$(1.3 \pm 0.4) \times 10^{-6}$
4	1921	-1.6 ± 1.0	1.6	$(3.7 \pm 4.6) \times 10^{-6}$	$(3.8 \pm 2.9) \times 10^{-7}$

TABLE 3
CENTRAL SZ TEMPERATURE DECREMENT AND INTEGRATED SZ SIGNAL FOR STACKED RADIO-QUIET LRGS, BINNED BY $L_{0.1r}$

Bin	N_{bin}	$\delta T_\circ$ μK	SNR	$Y_{200\bar\rho}$ arcmin^2	$Y_{200\bar\rho}\ d_A^2\ E(z)^{-2/3}$ Mpc^2
1	21	-28.2 ± 12.9	2.2	$(1.5 \pm 0.6) \times 10^{-4}$	$(1.1 \pm 0.5) \times 10^{-5}$
2	51	-10.6 ± 4.9	2.2	$(4.2 \pm 1.9) \times 10^{-5}$	$(3.7 \pm 1.7) \times 10^{-6}$
3	587	-6.1 ± 1.6	3.8	$(2.4 \pm 0.5) \times 10^{-5}$	$(1.7 \pm 0.5) \times 10^{-6}$
4	1732	-2.1 ± 1.0	2.1	$(7.8 \pm 6.9) \times 10^{-6}$	$(6.0 \pm 4.5) \times 10^{-7}$

NOTE. — The estimates for $Y_{200\bar\rho}$ are determined from $\delta T_\circ$ using microwave sky simulations (Sehgal et al. 2010a). They correspond to the halo masses $M_{200\bar\rho}$ derived from considering luminosity-dependent halo bias with $b_\star = 0.99$ (Table 1). The uncertainty in $\delta T_\circ$ is determined from bootstrap resampling. The central decrements and associated errors are measured in CMB temperature units.

significantly deeper than the stacked cluster at the center. For this stacked map, the pixel noise is around 5 μK, making the upper-edge cold spot a 5σ discrepancy. Examination of the 60 individual maps in the stack reveals that two maps contain decrements of < 150 μK at the location of the cold spot. In addition, 11 additional map sections show decrements larger than 50 μK at the approximate position of the upper-edge cold spot. This large number of decrements in similar positions relative to the LRG's in our catalog clearly is not a random phenomenon, but rather reflects the fact that galaxy clusters and groups are strongly clustered on the sky.

The errors on $\delta T_\circ$ presented in Tables 2 and 3 are computed for each bin using using bootstrap resampling. These errors measure the scatter in the properties of individual LRGs. Figure 3 shows the distribution of the mean of $\delta T_\circ$ for each of the four luminosity bins for the full and radio-quiet LRG samples. The distribution of each bin is well-approximated by a Gaussian distribution.

We use numerical simulations (Sehgal et al. 2010a) to relate the matched filter decrement to the integrated Y parameter (equation 1). In an unfiltered map, where the cluster is larger than the beam, $Y \simeq \delta T_0 \theta_c^2$, where θ_c is the angular size of the cluster. For a cluster of fixed physical size, $Y \propto \delta T_0 d_A^{-2}$ where d_A is the angular diameter distance to the cluster. Because the filter acts as a ∇^4 operator and removes large-scale signals in the map, the expected relationship for nearby clusters of fixed size is likely even steeper. However, rather than rely on this approximate scaling to weight the observations, we apply the same filter used on the ACT data to the simulated maps and compute $\delta T_{\circ,i}$ for each of the clusters. The simulated maps are first convolved with the ACT 148 GHz beam function. We then fit a simple model to the $Y - \delta T$ relation based on the angular diameter distance scaling:

$$Y_{est,i} = C\ \delta T_{\circ,i}\ z_i^{-\beta}, \qquad (5)$$

where z_i is the redshift and $Y_{est,i}$ serves as an estimate for the SZ signal within a disk of radius R_{200} for the ith

cluster. We define R_{200} as the radius within which the average density is equal to 200 times the mean matter density of the universe, $\bar\rho(z) = \rho_{crit}(z = 0)\Omega_m(1 + z)^3$. The relationship between Y and $\delta T_\circ$ is obtained for the mass range of each luminosity bin. Here the mass range is represented by the 1σ uncertainty in the central $M_{200\bar\rho}$ value, listed in Table 1. Due to systematic uncertainties in determining halo masses, we perform this analysis for three separate groups of mass bins (see Section 5.2).

The best fit values and corresponding errors for C and β are summarized in Table 4. The uncertainties reported are the formal 1σ errors computed using Levenberg–Marquardt least squares minimization. The correlation between Y_{est} and $Y_{200\bar\rho}$ is shown in Figure 4. Figure 5 shows the redshift dependence inherent in the relation between Y and $\delta T_\circ$, which motivated our model for Y_{est} in equation 5. Figures 4 and 5 show the results for the

TABLE 4
BEST-FIT PARAMETERS FOR THE $Y_{200\bar\rho}$ - $\delta T_\circ$ RELATION, OBTAINED FROM SIMULATION AND FITTED USING EQUATION 5

N_{bin}	Mass Range $10^{14} M_\odot$	C arcmin^2/K	β
		Halo Bias using $b_\star = 0.99$	
1103	$2.00 - 3.48$	1.37 ± 0.05	1.25 ± 0.03
1245	$1.65 - 1.91$	1.30 ± 0.05	1.18 ± 0.03
10321	$0.78 - 1.22$	1.08 ± 0.02	1.16 ± 0.02
15116	$0.29 - 0.57$	0.92 ± 0.05	1.14 ± 0.04
		Halo Bias using $b_\star = 1.30$	
382	$3.97 - 6.43$	1.47 ± 0.11	1.53 ± 0.07
291	$3.34 - 3.80$	1.42 ± 0.11	1.37 ± 0.07
2152	$1.76 - 2.58$	1.35 ± 0.04	1.24 ± 0.02
11274	$0.78 - 1.36$	1.09 ± 0.02	1.19 ± 0.02
		Weak Lensing Measurements	
3127	$1.72 - 3.46$	1.37 ± 0.03	1.30 ± 0.02
590	$1.62 - 2.48$	1.21 ± 0.07	1.24 ± 0.05
5408	$1.12 - 1.74$	1.16 ± 0.03	1.20 ± 0.02
11274	$0.60 - 1.02$	1.12 ± 0.03	1.08 ± 0.02

NOTE. — The three groups of mass ranges correspond to the three separate methods that are used to obtain halo masses for the LRG sample.

mass bins corresponding to $b_\star = 0.99$ in Table 1.

The values for $Y_{200\bar{\rho}}$ are found for the jth luminosity bin by summing over each LRG in the bin as follows,

$$Y_{200\bar{\rho},j} = \frac{\sum_i^{N_j} w_i \, C_j \, \delta T_{\circ,i} \, z_i^{-\beta_j}}{\sum_i^{N_j} w_i}, \qquad (6)$$

where $\delta T_{\circ,i}$ is the matched filter decrement recovered from the filtered ACT map and N_j is the number of LRGs in the jth luminosity bin. The weight w_i is related to the number of observations per pixel (see Section 4.2). Values for $Y_{200\bar{\rho}}$ are reported in Tables 2 and 3. These values correspond to the halo masses $M_{200\bar{\rho}}$ derived from considering luminosity-dependent halo bias with $b_\star = 0.99$. The error in $Y_{200\bar{\rho}}$ represents the uncertainty after propagating errors in $\delta T_\circ$, β, and C. These formal errors do not include the systematic uncertainties in our model which assumes that the gas profiles in Sehgal et al. (2010a) are accurate representations of the profiles for typical clusters.

Table 5 contains the matched filter decrements $\delta T_\circ$ obtained when stacking various subsets of the LRG sample. These subsets are designed to test various systematic aspects of the ACT data. The first four tests in Table 5 probe spatial discrepancies in the data by splitting the ACT map by right ascension and declination. We conclude that the signal is consistent throughout the map. We also examine variations in white noise amplitude by splitting the sample spatially by the number of observations N_{obs} per pixel. We determine the median value of N_{obs} per pixel and then stack subsets of LRGs that lie above and below this median value. We conclude that the signal is present for both high and low noise bins in the map.

Radio source contamination at 148 GHz appears to have a significant effect. Using the SDSS DR7 CAS, we identified LRGs with radio counterparts within $2''$ in the 1.4 GHz VLA FIRST Survey (Becker et al. 1995) and define these LRGs to be radio-loud. There are 290 radio-loud LRGs in the sample: 5/26 in Bin 1, 9/60 in Bin 2, 87/674 in Bin 3 and 189/1921 in Bin 4. Contamination by FIRST radio sources is potentially a concern for $\sim$ 10% of the LRG sample. We also compute the matched filter decrement for each bin excluding the 290 radio-loud LRGs. While the stacked radio-quiet LRGs show a significant ($> 4\sigma$) decrement, the radio-loud LRGs show clear evidence of radio source contamination. For the rest of this analysis, we report results for both the full and radio-quiet samples (see Tables 2 and 3).

5.2. *Estimates for Halo Mass* $M_{200\bar{\rho}}$

We take two different approaches to estimating the characteristic mass associated with halos of a given luminosity. The first approach is to use measurements of the halo bias, defined as the relation between the galaxy power spectrum $P_{gg}(k)$ and the linear matter power spectrum $P_{lin}(k)$,

$$P_{gg}(k) = b(k)^2 P_{lin}(k). \qquad (7)$$

On the large scales probed by the LRGs, the bias $b(k)$ is expected to be nearly scale-independent, but strongly dependent on luminosity. Tegmark et al. (2004) and

TABLE 5
MATCHED FILTER DECREMENT $\delta T_\circ$ FOR VARIOUS SYSTEMATIC
TESTS OF THE ACT DATA

Test	N_{bin}	$\delta T_\circ$ (μK)
Top	1295	-2.51 $\pm$ 1.16
Bottom	1386	-3.13 $\pm$ 1.10
Left	1406	-2.94 $\pm$ 0.99
Right	1275	-2.67 $\pm$ 1.35
High N_{obs}	1455	-2.15 $\pm$ 0.98
Low N_{obs}	1226	-3.85 $\pm$ 1.36
Radio-loud	290	3.27 $\pm$ 2.65
Radio-quiet	2391	-3.56 $\pm$ 0.84

NOTE. — The central decrements and associated errors are measured in CMB temperature units.

Zehavi et al. (2005) have fitted an empirical model for galaxy bias relative to the bias of $L_\star$ galaxies[1]:

$$\frac{b}{b_\star} = 0.85 + 0.15\frac{L}{L_\star} + 0.04(M_\star - M_{0.1r}), \qquad (8)$$

where $M_\star - 5\log_{10}(h) = -20.44$ (Blanton et al. 2003). For $h = 0.71$, $M_\star = -21.18$. Both $M_\star$ and $L_\star$ are in the $^{0.1}r$-band. Percival et al. (2007) have shown this model for relative bias to be a reasonable fit for the LRG sample considered in this paper.

There exists some uncertainty in the bias normalization factor, $b_\star$. Because this uncertainty can lead to large differences in halo mass, we will report mass estimates for two estimates of $b_\star$. The values used here are $b_\star = 0.99$ (Seljak et al. 2005) and $b_\star = 1.30$ (Reid et al. 2010).

Tinker et al. (2010) use numerical simulations to derive a relationship between halo mass and bias. They identify halos in their cosmological simulation using the spherical overdensity algorithm to define the halo mass:

$$M_\Delta = \frac{4}{3}\pi R_\Delta^3 \bar{\rho}(z)\Delta, \qquad (9)$$

where $\bar{\rho}(z)$ is the mean matter density of the universe at redshift z. This analysis will use $\Delta = 200$ and denote the halo mass by $M_{200\bar{\rho}}$. Tinker et al. (2010) estimate the relation between halo bias and $M_{200\bar{\rho}}$:

$$b(\nu) = 1 - A\frac{\nu^B}{\nu^B + \delta_c^B} + 0.183\nu^{1.5} + C\nu^{2.4}, \qquad (10)$$

where ν represents the "peak height" of the density field and is given by $\nu = \delta_c/\sigma(M_{200\bar{\rho}})$, δ_c is the critical density for collapse, and σ is the linear matter variance on the scale of the halo, i.e., $R = (3M/4\pi\bar{\rho})^{1/3}$. We use $z = 0$ to determine the matter variance σ in order to compare $b(\nu)$ to $b/b_\star$. In all calculations, we use $\delta_c = 1.686$. The parameters in equation 10 are:

$$A = 1.0 + 0.24y \, \exp[-(4/y)^4],$$
$$B = 0.44y - 0.88,$$
$$C = 0.019 + 0.107y + 0.19 \, \exp[-(4/y)^4],$$

where $y \equiv \log_{10}(200)$.

---

[1] $L_\star$ is defined with respect to the Schechter (1976) luminosity function for SDSS galaxies (see Blanton et al. 2003)

Using equations 8 and 10, we obtain estimates for halo mass for each luminosity bin. These values are given in Table 1. We determine mass from luminosity for each LRG in a given bin, and report $M_{200\bar{\rho}}$ as the mean of these values. The error on $M_{200\bar{\rho}}$ is estimated by the standard deviation of this mean, within each bin.

Stacked gravitational lensing measurements are an alternative method of estimating the halo mass-galaxy luminosity relation. Reyes et al. (2008) stack a sample of clusters from the SDSS maxBCG catalog (Koester et al. 2007a,b) and use weak gravitational lensing measurements to investigate the relation between cluster mass and various optical tracers. Since LRGs trace massive halos (Ho et al. 2009), the maxBCG catalog should represent a population similar to the LRG sample used in this work. Noting that Reyes et al. (2008) bin their cluster sample by $^{0.25}r$-band BCG luminosity, we compute the $^{0.25}r$-band luminosity for each LRG and determine halo masses using equation 15(c) of Reyes et al. (2008). Table 1 lists these mass values, as well as those obtained from analysis of halo bias. The range in these three masses represents some of the systematic uncertainties in relating LRG luminosity to halo mass.

6. SZ–MASS SCALING RELATION

In this section, we present an analysis of the correlation between cluster mass and integrated SZ signal. Specifically, we compare $Y_{200\bar{\rho}}$ to the cluster mass $M_{200\bar{\rho}}$ for the luminosity/mass bins in Table 1. For self-similar evolution, analytical theory (e.g., Komatsu & Seljak 2001) predicts a simple relationship between the integrated Y parameter and cluster mass. Assuming virial and hydrostatic equilibrium, the cluster gas temperature can be related to M and $E(z)$ through (e.g., Bryan & Norman 1998)

$$T \propto M^{2/3} E(z)^{2/3}, \qquad (11)$$

where for a flat ΛCDM cosmology

$$E(z) = [\Omega_m(1+z)^3 + \Omega_\Lambda]^{1/2}. \qquad (12)$$

Under assumptions of an isothermal ICM,

$$Y_{200\bar{\rho}} \propto d_A^{-2}(z) \, M_{200\bar{\rho}} \, T, \qquad (13)$$

where $d_A(z)$ is the angular diameter distance to the cluster. Combining the above equations leads to the expectation for self-similar evolution

$$Y_{200\bar{\rho}} \propto d_A^{-2} \, M_{200\bar{\rho}}^{5/3} \, E(z)^{2/3}. \qquad (14)$$

Anticipating this scaling relation, we have plotted $Y_{200\bar{\rho}}d_A^2 E^{-2/3}$ against the halo masses $M_{200\bar{\rho}}$ in Figure 6. We show the three separate halo mass estimates for a given bin to illustrate the uncertainties that exist in converting LRG luminosity to halo mass. The results for both the full (Table 2) and radio-quiet (Table 3) LRG samples are shown in this figure. The expected $Y - M$ relation as determined from numerical simulations is shown as a solid line and is given by:

$$Y_{200\bar{\rho}}d_A^2 E^{-2/3^i} = 10^\gamma \left(\frac{M_{200\bar{\rho}}}{10^{14}M_\odot} \right)^\alpha, \qquad (15)$$

where the best fit parameters are $\alpha = 1.76$ and $\gamma = -5.74$. These parameters are obtained from the microwave sky simulations described in Sehgal et al. (2010a), where halos of mass $M_{200\bar{\rho}} > 2.82 \times 10^{13} \, M_\odot$ have been considered. Furthermore, the redshift range of the clusters is restricted to $0.15 < z < 0.50$, and a cylinder definition of Y is used. Tables 2 and 3 give the values of $Y_{200\bar{\rho}}d_A^2 E^{-2/3}$ and associated errors for the full and radio-quiet samples. These estimates are determined using a summation similar to equation 6, where the redshift of each LRG is used to calculate $d_A(z)$ and $E(z)$. They correspond to the masses derived from considering luminosity-dependent halo bias with $b_\star = 0.99$.

7. DISCUSSION

By stacking LRGs, we have obtained estimates of the SZ signal as function of mass for groups and clusters. The stacking analysis enabled the detection of an SZ signal for clusters with masses as low as $\sim 3 \times 10^{13} M_\odot$, and the amplitude of the signal is consistent with our theoretical estimates. Using masses derived from halo bias, the slope of the inferred $Y - M$ relation is consistent with expected results. However, when using stacked weak gravitational lensing measurements to estimate the halo masses, the slope appears to be slightly steeper than the predicted value. Agreement with the overall normalization of the $Y - M$ relation depends upon our method of estimating the halo mass from luminosity, which is subject to large uncertainties.

We have detected radio source contamination of the SZ signal at 148 GHz for $\sim 10\%$ of the LRGs. This suggests the possibility that infrared sources may also have a contamination effect on our measurement of the SZ signal for the LRG sample. Mittal et al. (2009) found that the likelihood for a cluster to host a radio-loud BCG is a strong function of the central cooling time, in the sense that all strong cool-core clusters contain a radio-loud BCG. Thus, our radio-quiet LRG sample may select against clusters/groups with a strong cool core. However, in his analysis of a large sample of clusters and groups at $z \leq 0.1$, Sun (2009) did not find any radio-loud central galaxies in groups with strong cool cores. If this is true at higher redshift, our radio-quiet LRG sample would be representative of all groups and clusters.

Another significant source of systematic error is the conversion of the measured decrement $\delta T_\circ$ to the integrated SZ signal Y. Our approach relies on simulations from Sehgal et al. (2010a) to estimate the effect of the matched filter on smoothing the cluster signal. Future work will explore this uncertainty by using simulations with different models of cluster gas physics.

This work is a first application of a novel stacking method for searching for the SZ effect in clusters by applying a full optical catalog to a CMB survey. Both data sets are rich and have multiple components, with possible correlations between them, that we will be sorting out as the data and our understanding improves. Future ACT analysis of the SZ effect from galaxy clusters will include data from multiple seasons and millimeter bands and will include the new SDSS DR8 analysis. Our work shows that the presence of radio sources coincident with LRGs affects the 148 GHz SZ signal. Our maps show

REFERENCES

Abazajian, K., et al. 2003, AJ, 126, 2081
—. 2004, AJ, 128, 502
—. 2005, AJ, 129, 1755
Abazajian, K. N., et al. 2009, ApJS, 182, 543
Adelman-McCarthy, J. K., et al. 2006, ApJS, 162, 38
—. 2007, ApJS, 172, 634
—. 2008, ApJS, 175, 297
Andersson, K., et al. 2010, ArXiv e-prints
Bartlett, J. G., Chamballu, A., Melin, J., Arnaud, M., & Members of the Planck Working Group 5. 2008, Astronomische Nachrichten, 329, 147
Becker, R. H., White, R. L., & Helfand, D. J. 1995, ApJ, 450, 559
Benson, B. A., Church, S. E., Ade, P. A. R., Bock, J. J., Ganga, K. M., Henson, C. N., & Thompson, K. L. 2004, ApJ, 617, 829
Birkinshaw, M. 1999, Phys. Rep., 310, 97
Blanton, M. R., & Roweis, S. 2007, AJ, 133, 734
Blanton, M. R., et al. 2003, ApJ, 592, 819
Bonamente, M., Joy, M., LaRoque, S. J., Carlstrom, J. E., Nagai, D., & Marrone, D. P. 2008, ApJ, 675, 106
Bryan, G. L., & Norman, M. L. 1998, ApJ, 495, 80
Carlstrom, J. E., Holder, G. P., & Reese, E. D. 2002, ARA&A, 40, 643
Carlstrom, J. E., et al. 2009, ArXiv e-prints
Das, S., et al. 2010, ArXiv e-prints
Dunkley, J., et al. 2010, ArXiv e-prints
Eisenstein, D. J., et al. 2001, AJ, 122, 2267
Fowler, J. W., et al. 2007, Appl. Opt., 46, 3444
—. 2010, ApJ, 722, 1148
Fukugita, M., Ichikawa, T., Gunn, J. E., Doi, M., Shimasaku, K., & Schneider, D. P. 1996, AJ, 111, 1748
Gunn, J. E., et al. 1998, AJ, 116, 3040
Haehnelt, M. G., & Tegmark, M. 1996, MNRAS, 279, 545
Hajian, A., et al. 2010, ArXiv e-prints
Herranz, D., Sanz, J. L., Barreiro, R. B., Hobson, M., Martinez-Gonzalez, E., & Diego, J. M. 2002, in Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, Vol. 4847, Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, ed. J.-L. Starck & F. D. Murtagh, 50–61
High, F. W., et al. 2010, ApJ, 723, 1736
Hincks, A. D., et al. 2010, ApJS, 191, 423
Ho, S., Lin, Y., Spergel, D., & Hirata, C. M. 2009, ApJ, 697, 1358
Huang, C., et al. 2010, ApJ, 716, 758
Kazin, E. A., et al. 2010, ApJ, 710, 1444
Koester, B. P., et al. 2007a, ApJ, 660, 239
—. 2007b, ApJ, 660, 221

Komatsu, E., & Seljak, U. 2001, MNRAS, 327, 1353
Komatsu, E., et al. 2009, ApJS, 180, 330
Lupton, R., Gunn, J. E., Ivezić, Z., Knapp, G. R., & Kent, S. 2001, in Astronomical Society of the Pacific Conference Series, Vol. 238, Astronomical Data Analysis Software and Systems X, ed. F. R. Harnden Jr., F. A. Primini, & H. E. Payne, 269–+
Mantz, A., Allen, S. W., & Rapetti, D. 2010, MNRAS, 406, 1805
Marriage, T. A., et al. 2010a, ArXiv e-prints
—. 2010b, ArXiv e-prints
Melin, J., Bartlett, J. G., & Delabrouille, J. 2006, A&A, 459, 341
Melin, J., Bartlett, J. G., Delabrouille, J., Arnaud, M., Piffaretti, R., & Pratt, G. W. 2010, ArXiv e-prints
Menanteau, F., et al. 2010, ApJ, 723, 1523
Mittal, R., Hudson, D. S., Reiprich, T. H., & Clarke, T. 2009, A&A, 501, 835
Percival, W. J., et al. 2007, ApJ, 657, 645
Plagge, T., et al. 2010, ApJ, 716, 1118
Reid, B. A., & Spergel, D. N. 2009, ApJ, 698, 143
Reid, B. A., et al. 2010, MNRAS, 404, 60
Rephaeli, Y. 1995, ARA&A, 33, 541
Reyes, R., Mandelbaum, R., Hirata, C., Bahcall, N., & Seljak, U. 2008, MNRAS, 390, 1157
Richards, G. T., et al. 2002, AJ, 123, 2945
Rozo, E., et al. 2010, ApJ, 708, 645
Schechter, P. 1976, ApJ, 203, 297
Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, ApJ, 500, 525
Sehgal, N., Bode, P., Das, S., Hernandez-Monteagudo, C., Huffenberger, K., Lin, Y., Ostriker, J. P., & Trac, H. 2010a, ApJ, 709, 920
Sehgal, N., et al. 2010b, ArXiv e-prints
Seljak, U., et al. 2005, Phys. Rev. D, 71, 043511
Staniszewski, Z., et al. 2009, ApJ, 701, 32
Strauss, M. A., et al. 2002, AJ, 124, 1810
Sun, M. 2009, ApJ, 704, 1586
Sunyaev, R. A., & Zel'dovich, Y. B. 1970, Ap&SS, 7, 20
Swetz, D. S., et al. 2010, ArXiv e-prints
Tegmark, M., et al. 2004, ApJ, 606, 702
Tinker, J. L., Robertson, B. E., Kravtsov, A. V., Klypin, A., Warren, M. S., Yepes, G., & Gottlöber, S. 2010, ApJ, 724, 878
Vanderlinde, K., et al. 2010, ApJ, 722, 1180
Vikhlinin, A., et al. 2009, ApJ, 692, 1060
Zehavi, I., et al. 2005, ApJ, 630, 1
Zel'dovich, Y. B., & Sunyaev, R. A. 1969, Ap&SS, 4, 301
Zheng, Z., Zehavi, I., Eisenstein, D. J., Weinberg, D. H., & Jing, Y. P. 2009, ApJ, 707, 554
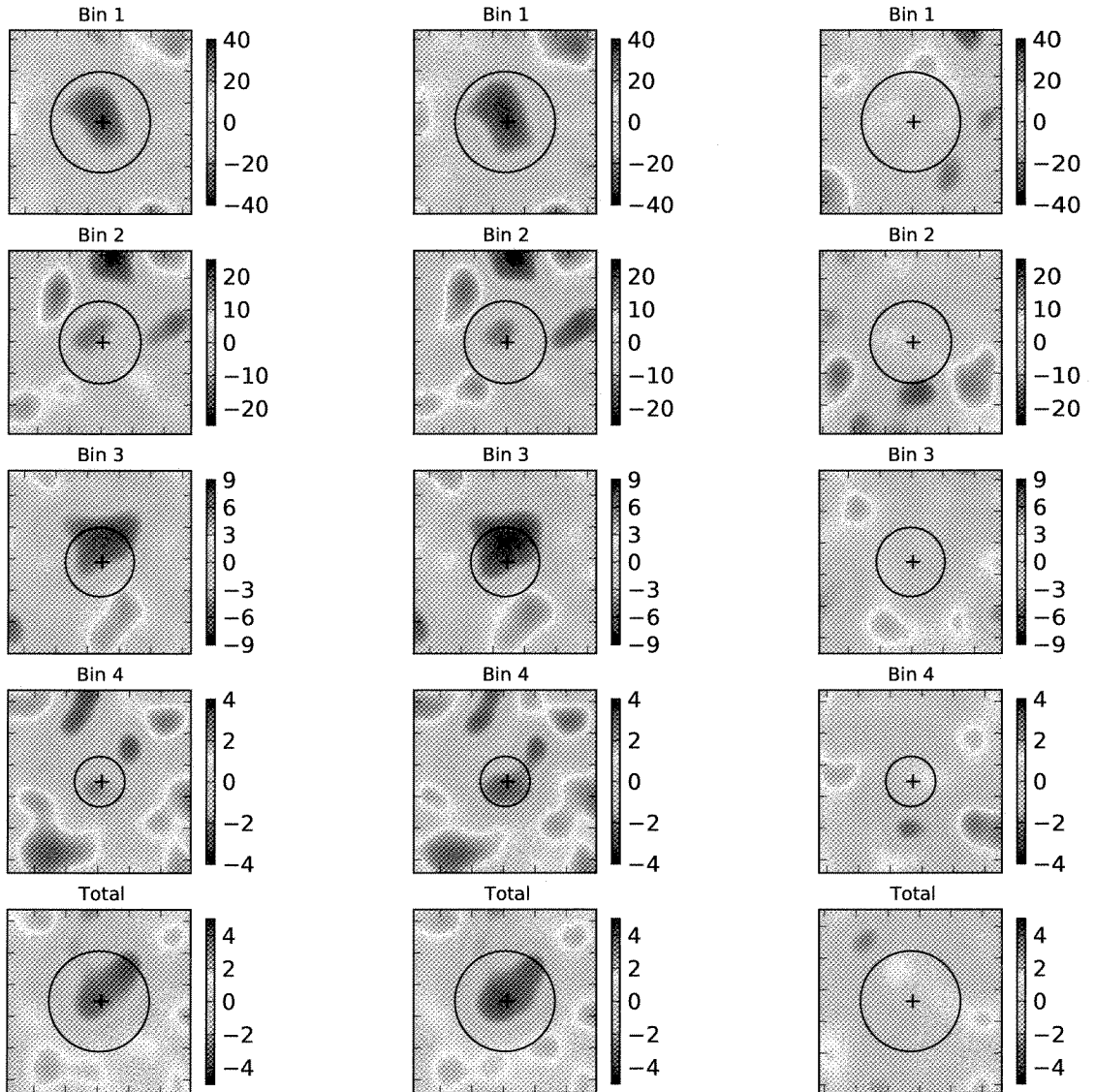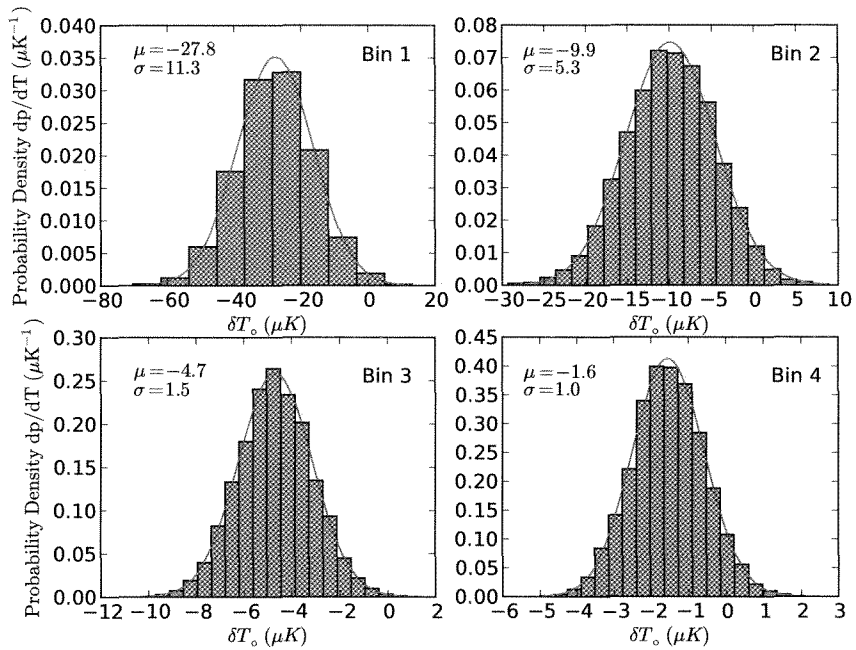
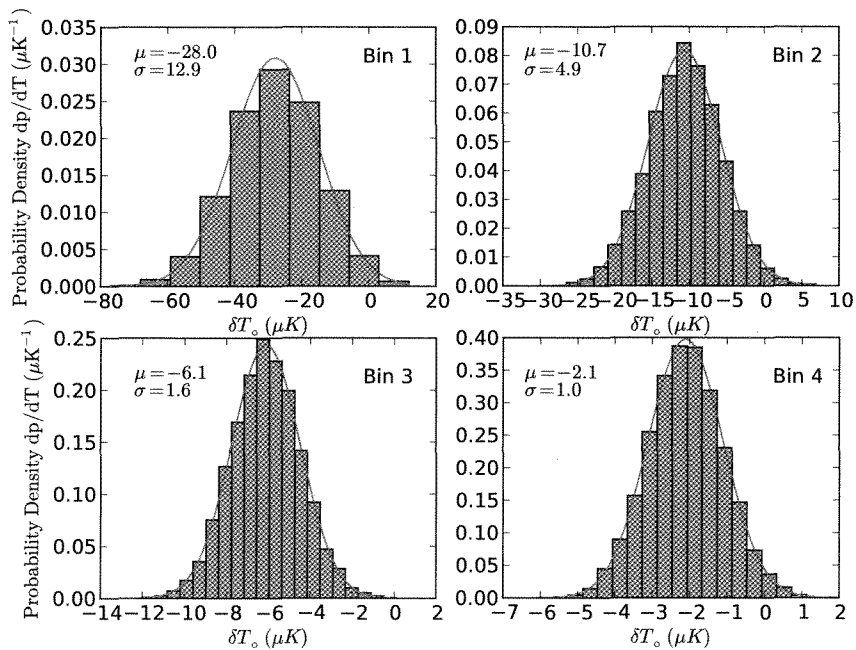(a) Full LRG Sample   (b) Radio-quiet LRG Sample   (c) Random Sample

FIG. 2.— The stacked filtered temperature maps for each luminosity bin for the (a) full LRG sample, (b) radio-quiet LRG sample and (c) a randomly selected sample. Each map is $7' \times 7'$ and is in CMB temperature units (μK). The values for $\delta T_\circ$ and the number of maps in each stack are given in Tables 2 and 3 for (a) and (b), respectively. The radio-quiet sample excludes the 10% of the full LRG sample that has been defined as radio-loud (see Section 5.1). The stacked random maps serve as a null test of the data, as the SZ signal should be consistent with zero for these maps. The number of maps in each stack for the random sample is equal to that of the full sample. The circles in each submap mark the characteristic size of a cluster for a given luminosity/mass bin, in terms of θ_{200}. The values for θ_{200} have been determined from simulation. The crosses mark the center of each map.

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(a) Full LRG Sample

(b) Radio-quiet LRG Sample

FIG. 3.— The distribution of the mean of $\delta T_{\circ}$ for each luminosity bin in the full and radio-quiet LRG samples, determined from bootstrap resampling. The mean μ and standard deviation σ obtained from the resampled distribution are shown for each bin. The distribution of the mean is well-approximated by a Gaussian function, which is shown in red for each panel. The bootstrap errors σ are reported in Tables 2 and 3 and are used throughout this work as the error estimate for $\delta T_{\circ}$.
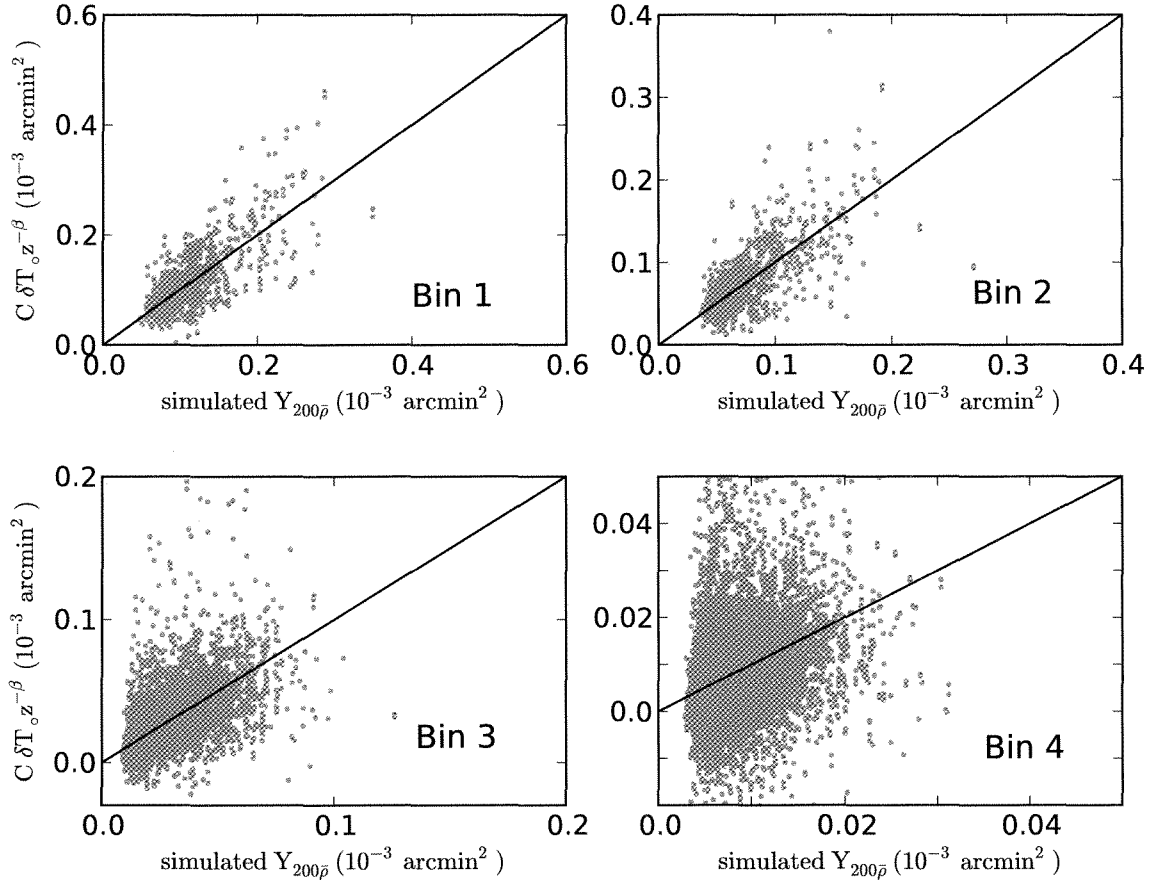
FIG. 4.— The correlation between $\delta T_\circ$ and $Y_{200\bar{\rho}}$ for each luminosity/mass bin, as determined from simulations described in Sehgal et al. (2010a). The simulation values for $Y_{200\bar{\rho}}$ of the ith cluster are fit to the model $Y_{est,i} = C \ \delta T_{\circ,i} \ z_i^{-\beta}$. The four subplots correspond to the luminosity bins used in the LRG stacking analysis. The mass range for each bin corresponds to the mass estimates determined from using halo bias, with $b_\star = 0.99$ (Table 1).
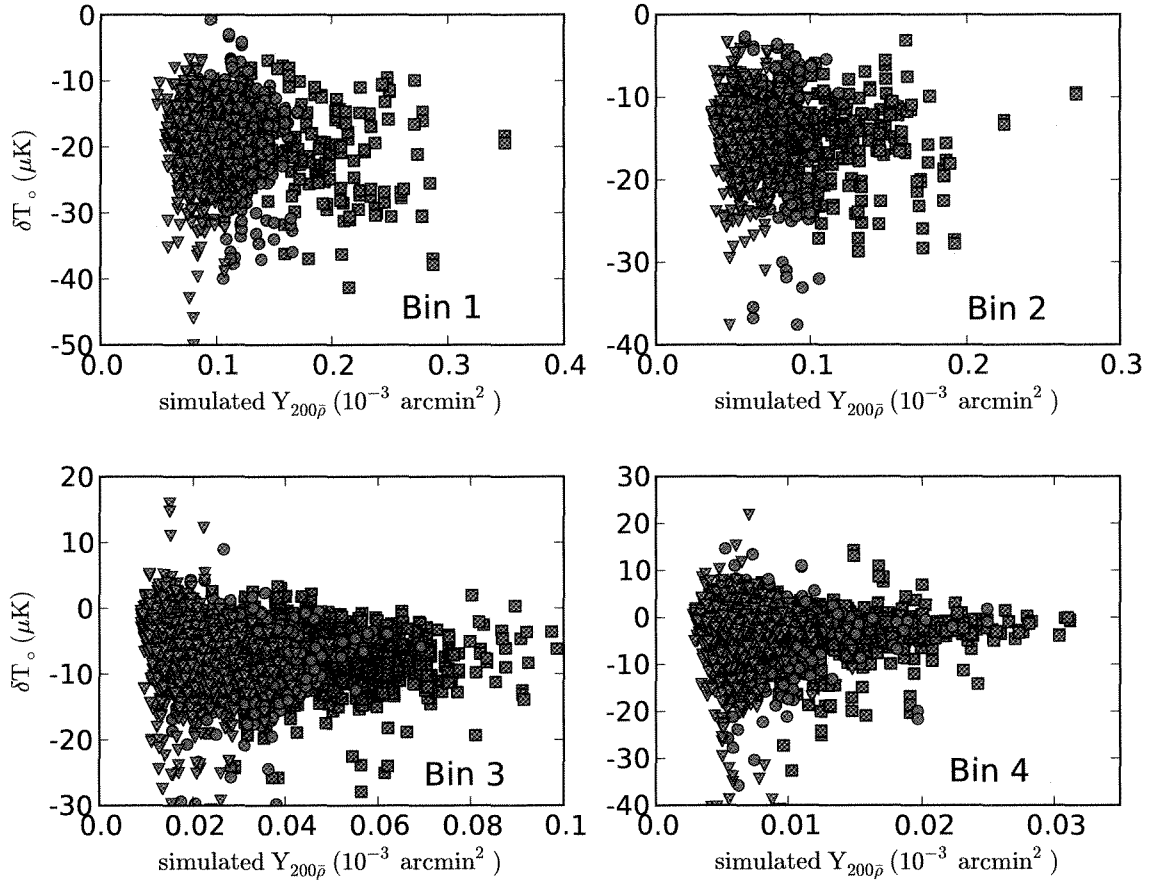
FIG. 5.— The redshift dependence of the relation between $\delta T_\circ$ and $Y_{200\bar{\rho}}$ for each luminosity/mass bin, as determined from simulations described in Sehgal et al. (2010a). The redshift ranges are represented as follows: $z < 0.26$ (blue squares), $0.26 < z < 0.37$ (red circles), and $z > 0.37$ (green triangles). The four subplots correspond to the luminosity bins used in the LRG stacking analysis. The mass range for each bin corresponds to the mass estimates determined from using halo bias, with $b_\star = 0.99$ (Table 1).
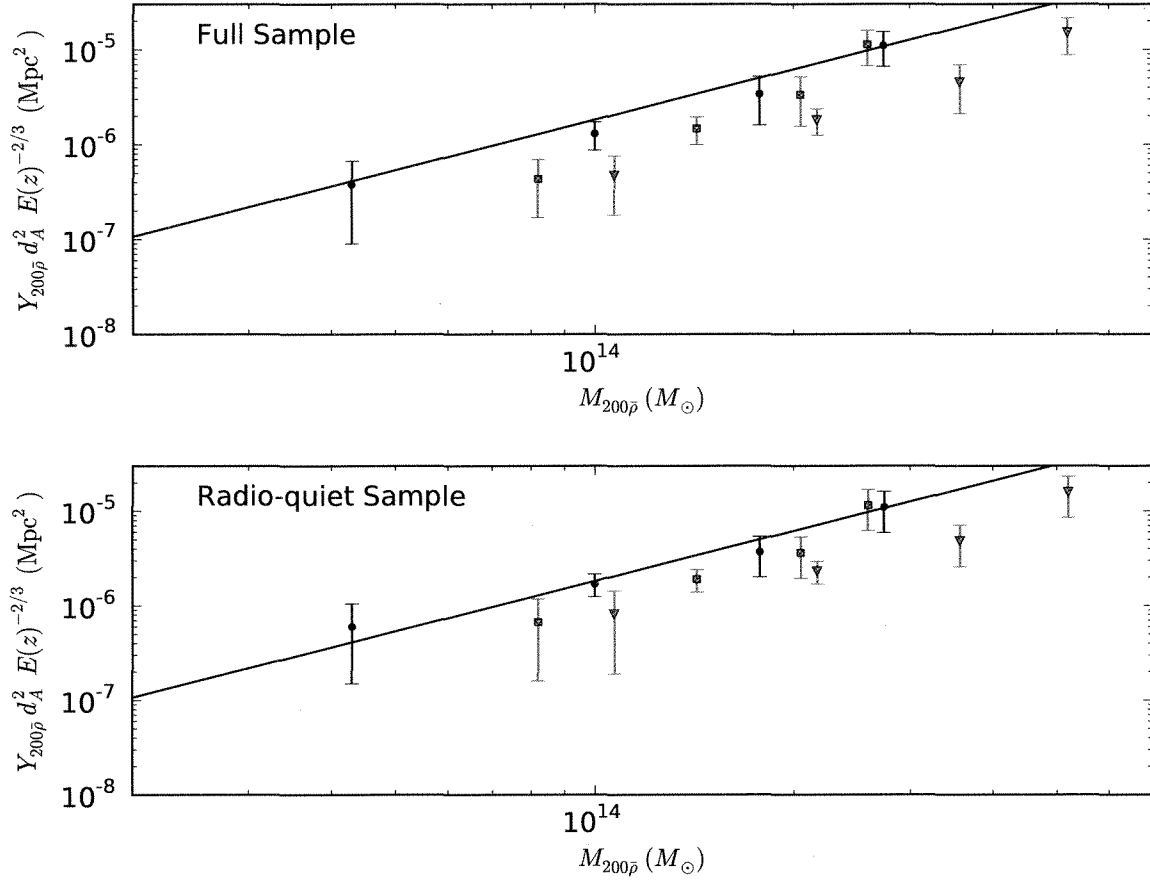
FIG. 6.— SZ – Mass Correlation. $Y_{200\bar{\rho}}$ estimates are plotted against cluster masses for each luminosity bin for both the full (Table 2) and radio-quiet (Table 3) LRG samples . The three sets of points shown use three separate mass estimates for the four LRG luminosity bins. The black circles correspond to mass estimates derived from analysis of halo bias using $b_\star = 0.99$ (Seljak et al. 2005). The red triangles correspond to mass estimates derived from analysis of halo bias using $b_\star = 1.30$ (Reid et al. 2010). Finally, the blue squares correspond to mass estimates derived from weak gravitational lensing measurements (Reyes et al. 2008). The range in these masses represents some of the systematic uncertainties in converting LRG luminosity to halo mass. The solid line shows the expected model for the $Y - M$ relation (equation 15), as determined from the microwave sky simulations of Sehgal et al. (2010a).