

# Twelve and a Half Years of Observations of Centaurus A with *RXTE*

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## ABSTRACT

The *Rossi X-ray Timing Explorer* has observed the nearest radio galaxy, Centaurus A, in 13 intervals from 1996 August to 2009 February over the 3 – 200 keV band. Spectra accumulated over the 13 intervals were **well** described with an absorbed power law and an iron line. Cut-off power laws and Compton reflection from cold matter did not provide a better description. **For the 2009 January observation we set** a lower limit on the cutoff energy at over 2 MeV. The power spectral density function was generated from *RXTE*/ASM and PCA data as well as an *XMM-Newton* long look, and clear evidence for a break at  $18_{-7}^{+18}$  days (**68% conf.**) was seen. Given Cen A’s high black hole mass and very low value of  $L_X/L_{\text{Edd}}$ , the break was a factor of  $17 \pm 9$  times higher than the break frequency predicted by the McHardy and coworkers relation, which was empirically derived for **a sample of objects, which are** radio-quiet **and** accreting at relatively high values of  $L_{\text{bol}}/L_{\text{Edd}}$ . We have interpreted our observations in the context of a clumpy molecular torus. The variability characteristics and the broadband spectral energy distribution, when compared to Seyferts, imply that the bright hard X-ray continuum emission may originate at the base of the jet, yet from behind the absorbing line of sight material, in contrast to what is commonly observed from blazars.

*Subject headings:* radio galaxies: general — radio galaxies: individual (Cen A)

## 1. Introduction

The prototype **Fanaroff-Riley Class I (FR I; Fanaroff & Riley 1974)** low luminosity radio source Centaurus A (hereafter Cen A) played a significant role in the early understanding of the nature of active galaxies and jet emission. Cen A was found to “correspond within limits of experimental error” to the “extra-galactic nebula” NGC 5128 by Bolton, Stanley, & Slee (1949). Early radio interferometry by Mills (1952) established that Cen A was an extended source and therefore not a stellar object. In addition, Mills found it to be nearly  $2^\circ$  in diameter in radio with “a strong concentration near its centre”. A year later Mills (1953) refined the sizes to  $1.5^\circ$  for the extended flux and  $5'$  for the compact core. Arguments by Baade & Minkowski (1954) demonstrated that Cen A was extragalactic, and Humason, Mayall, & Sandage (1956) revealed an uncorrected redshift close to the present day value ( $z = 0.001825$ ). The giant radio lobes were resolved by Wade (1959) after removal of the bright compact source coincident with NGC 5128. The synchrotron nature of the radio emission was put forth by Shklovsky and collaborators (see, e.g., Burbidge & Burbidge 1957, for references).

The Centaurus region of the southern sky was scanned in the late 1960s by rocket and balloon X-ray payloads without detecting Cen A (Cooke et al. 1967; Harries et al. 1967; Lewin, Clark, & Smith 1968). The 1969 rocket flight of Bowyer et al. (1970) was the first claimed detection, and was followed by many rocket, balloon and satellite observations (see Stark, Davison, & Culhane 1976, for a list of early results). A power law continuum with photon index of  $\Gamma \sim 1.8$  extending beyond 200 keV was soon established with low energy absorption due to intervening matter having an equivalent hydrogen column density,  $N_{\text{H}}$ , along the line of sight near  $10^{23}$  atoms  $\text{cm}^{-2}$  (e.g., Baity et al. 1981). *CGRO* observations of Cen A extended the flux detections to 1 MeV (Kinzer et al. 1995), and to greater than 100 MeV (Steinle et al. 1998). Recently, emission from the core of Cen A (Abdo et al. 2010b) and its extended emission (Abdo et al. 2010a) have been detected by the *Fermi*/Large Area Telescope (LAT) at GeV energies, and by *H.E.S.S.* at TeV energies

(Aharonian et al. 2009), further extending the detected emissions from this remarkable object over the entire electromagnetic spectrum.

By virtue of being the nearest radio galaxy and one of the few bright active galaxies in the 2–10 keV X-ray band, detailed X-ray observations of Cen A revealed much about the nuclei of low power, radio-loud active galaxies in general (see Israel 1998, for a review). With present day imaging instruments, we have a remarkably detailed view of Cen A. *Chandra*'s spectacular images have revealed details for a wealth of features: jet structure, knots in the jets, discrete sources, and shocks in the interstellar medium (e.g., Kraft et al. 2000, 2001, 2003). Evans et al. (2004) used *Chandra* and *XMM-Newton* observations to conclude that the iron line at 6.4 keV originated in cool material that either surrounds the black hole or forms part of a torus, most of which is out of the line of sight. In either case the iron emitting region was distant from the central source.

While comparison of the blazar Cen A to Seyfert galaxies is not strictly proper due to jet dominated blazar spectra versus disk dominated Seyfert spectra, Cen A displays **at least one** aspect of disk dominated spectra, i.e. iron line emission from cold material. The jet-like emission comes from a jet whose primary beam is oriented far from the observer's line of sight and thus some Seyfert-like aspects of the spectra become detectable. One must keep this in mind when comparing Cen A results to those of Seyfert AGN and to those of blazars.

Infrared observations from the ground with adaptive optics and from *Hubble Space Telescope* yielded estimates of the mass of the supermassive black hole from gas kinematics and from stellar dynamics (e.g., Silge et al. 2005; Marconi et al. 2006). Integral-field observations in the near infrared with the SINFONI spectrograph with adaptive optics on the VLT gave  $M_{\text{BH}}=(5.5\pm 3.0)\times 10^7 M_{\odot}$  based upon stellar kinematics and  $(4.5^{+1.7}_{-1.0})\times 10^7 M_{\odot}$  from  $\text{H}_2$  kinematics (Cappellari et al. 2009; Neumayer et al. 2010). Krajnović, Sharp, & Thatte (2007), using the CIRCASS spectrograph on Gemini South, found  $M_{\text{BH}}=(8.25^{+2.25}_{-4.25})\times 10^7 M_{\odot}$  based on gas dynamics. For a distance of 3.42 Mpc (Ferrarese et

al. 2007), 3.5 Mpc (Hui et al. 1993), or 3.8 Mpc (Rejkuba 2004; Harris, Rejkuba, & Harris 2010), this mass must be contained within  $\simeq 0.6$  pc (Marconi et al. 2006). **We adopt the values of 3.8 Mpc and  $6 \times 10^7 M_{\odot}$  for the following analyses and discussions.**

Radio observations of the inner tens of parsecs reveal an unresolved core with a jet and counter jet (Horiuchi et al. 2006) with the inner  $\pm 20$  mas being resolved into multiple components. TANAMI radio observations (Ojha et al. 2010) find that the core region has an inverted spectrum indicative of on-going synchrotron and possible synchrotron self-Compton (SSC) or free-free absorption processes (Müller et al. 2010). While not unexpected, this would support a model of synchrotron self-Compton for producing the hard X-ray flux observed (Chiaberge, Capetti, & Celotti 2001).

With its broad band X-ray coverage, highly maneuverable spacecraft, and flexible scheduling, the *Rossi X-ray Timing Explorer (RXTE)* is the premier mission for monitoring and multiple observations of X-ray sources on timescales from milliseconds to years. Multi-timescale monitoring campaigns to probe X-ray variability in active galactic nuclei (AGN) on time scales from hours to years began with the *EXOSAT* era in the 1980s (e.g., Lawrence et al. 1987; Lawrence & Papadakis 1993; Green et al. 1993), and has continued to present day with *RXTE* (e.g., Edelson & Nandra 1999; Markowitz et al. 2003; McHardy et al. 2006). The resulting broadband Power Spectral Density functions (PSDs), derived mainly for radio quiet “normal” broad line Seyfert 1s and narrow-line Seyfert 1s, have yielded evidence for breaks at temporal frequencies  $f_b$  in the range  $\sim 10^{-6}$  to  $10^{-3}$  Hz with PSD power-law slopes breaking from about  $-2$  to about  $-1$  above and below  $f_b$ , respectively. McHardy et al. (2006) and references therein have shown that the temporal variability resulting from accretion onto black holes, as characterized by  $f_b$ , scales inversely with black hole mass and directly with bolometric luminosity. Comparatively less is known about the X-ray variability properties of non-Seyferts, including jet-dominated blazars and radio galaxies such as Cen A. Preliminary X-ray structure functions and PSDs for blazars, though frequently based on long-term light curves which are gap-dominated, have revealed

breaks which, like Seyferts, correspond to timescales of a few days to a couple of weeks (Kataoka et al. 2001, 2002; Marscher et al. 2004).

In this paper we present spectral and timing results on Cen A derived from 13 observations spanning 1996 to 2009 made with *RXTE*. Our goals include characterizing broadband X-ray spectral variability, constraining the geometry of the circumnuclear accreting material by studying both the line of sight absorbing gas and the Fe  $K\alpha$  line-emitting gas, and quantifying X-ray variability across a range of time scales in order to better understand Cen A for comparison with other AGN. The remainder of this paper is structured as follows: Section 2 gives the details of the observations, Section 3 describes data selection and analysis techniques, Section 4 gives the results of spectral and temporal analyses, including a first-ever broadband PSD (covering below  $\sim 10^{-5}$  Hz) for Cen A, while Section 5 discusses the results of these observations of Cen A, and finishes with our conclusions.

## 2. Observations

The flux history of Cen A over the duration of the *RXTE* mission from 1996 through mid 2010, as measured in the 5–12.1 keV band by the *RXTE*/All-Sky Monitor (Fig. 1), reveals variations of a factor of 3 or more. Thirteen observing campaigns totaling 129 separate pointings were executed using *RXTE*'s pointed-mode instruments, the Proportional Counter Array (PCA) and the High Energy X-ray Timing Experiment (HEXTE), as indicated in Fig. 1. Each campaign consisted of 1 to 22 individual pointings, and spanned durations of less than a day to several days. The 13 campaigns could be divided into observations on 51 separate days (from which 47 daily spectra were produced; see §4.2.2), and adjacent campaigns were separated by less than a month to years. We thus have the ability to probe spectral and temporal variability in Cen A on a wide range of time scales, including hours, individual days, and weeks to years.

Table 1 gives the livetimes for the 13 separate observing intervals of Cen A from 1996 August to 2009 February. The counting rates are for 3–60 keV from the top layer of PCA detector PCU2 and 15–200 keV for HEXTE. Since the on-source/off-source rocking of HEXTE cluster A was terminated on 2006 July 13, the livetimes and rates are given for the combined clusters up to that date as well as just cluster B throughout the mission. See the next Section for a description of the PCA and HEXTE.

### 3. Data Analysis

Version v6.7 of the HEASOFT software package release<sup>1</sup> with updated PCA response v11.7 (2009 May 11) and background estimation files Faint/L7 and Sky\_VLE dated 2005 Nov. 28 were used throughout, as was XSPEC 12.5.1n (Arnaud 1996). All errors presented are 90% confidence intervals, with the exceptions of iron line equivalent widths, which have 99% uncertainties, and the counting rates in Table 1 are 68% errors.

#### 3.1. Proportional Counter Array

For the PCA (Jahoda et al. 2006), only data from the top layer of the second Proportional Counter Unit (PCU) — PCU2 — were used in the present analysis, since its calibration was the most refined of the 5 PCU counters (Jahoda et al. 2006), and since it was the one PCU in common for all of the Cen A observations. The PCU2 data were accumulated under the conditions that the source be greater than  $10^\circ$  above the Earth’s limb, the pointing direction be within  $0^\circ.01$  of the source position, the satellite was more than 30 minutes from the last onset of a South Atlantic Anomaly (SAA) passage, and the veto rate indicative of precipitating electrons was less than  $0.1^2$ . The PCU2 Standard

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<sup>1</sup><http://heasarc.gsfc.nasa.gov/ftools/>

<sup>2</sup>[http://heasarc.gsfc.nasa.gov/docs/xte/recipes/pca\\_event\\_spectra.html#reduction](http://heasarc.gsfc.nasa.gov/docs/xte/recipes/pca_event_spectra.html#reduction)

Data counts histograms were rebinned into wider energy bins above channel 60 by 2, above channel 80 by 4, and above channel 100 by 29 in all analyses. PCU2 data from the full 3–60 keV energy range were subject to fitting for this paper, and no systematic errors were added to the PCU2 data.

The Faint background model was used for all but the 2009 January/February observations, for which the SkyVLE background model was used. The choice of background models was determined by whether or not the PCU2 toplayer rate exceeded  $50 \text{ counts s}^{-1}$ , and only the 2009 January and February rates exceeded this value. The PCU2 background models were calculated from particle rates from the detector based upon numerous blank sky pointings throughout the *RXTE* mission (Jahoda et al. 2006), and as such, individual observations were susceptible to minor differences in normalization between the model background and reality. Similarly, the deadtime correction was based upon detector rates and modeled appropriately. In order to adjust spectral analyses for small deviations of the background and deadtime models from reality, the background model counts histogram was included in the XSPEC fitting as a correction file (see § 3.3 for a description of this procedure).

### 3.2. High Energy X-ray Timing Experiment

The HEXTE (Rothschild et al. 1998) data from on-source and two off-source positions were collected from each cluster separately until the cluster A rocking was discontinued on 2006 July 13. After that the cluster B data were still collected from both off-source positions, but no off-source data were collected from cluster A. The same elevation, pointing accuracy, and time since the SAA criteria were used in the selection of good data for HEXTE. The HEASARC tool HEXTERock was used to determine if a known source would be in any of the HEXTE off-source positions. None were. Consequently, the two background spectral accumulations, when available for cluster A and always for cluster B, were combined to form a single background spectrum for each cluster. The on-source



accumulations for both clusters were then combined as well as the appropriate background files. Combining these data sets directly was possible due to the HEXTE automatic and continuous gain control system that allowed for the setting of all detector gains to be the same. This resulted in a single on-source and single background file for each Cen A observation for which both clusters were rocking. In those instances where cluster A was not rocking, only cluster B data were used. No systematic errors were added to the HEXTE data.

HEXTE background files were accumulated in real-time during an observation as a result of the HEXTE cluster rocking program, and were not a product of a modeling program. The deadtime calculated for each data segment was, however, subject to a model based upon particle event rates, and as such could produce imperfect background subtraction. The same technique for dealing with this effect as for PCU2 was used, and the correction amounted to less than 1% in all cases (Figure 2). The HEXTE data were rebinned by 5 above channel 60, by 10 above channel 160, and by 20 above channel 210. HEXTE data from 18–200 keV were used in this paper.

### 3.3. Simultaneous PCA and HEXTE Spectral Fitting

Simultaneous fitting of the PCU2 and HEXTE data was performed on both daily and the 13 observation interval data where the HEXTE data had sufficient statistical quality to affect the result. This provided the best-fit *RXTE* broad band 3–200 keV spectral parameters for Cen A with overlapping PCU2 and HEXTE data in the 18–60 keV band. A constant multiplying the HEXTE model was fitted to correct for the relative normalization of HEXTE with respect to PCU2. The fitted value of this constant was about 80% for the observations. This reflected the differences in assumed Crab normalization between PCA and HEXTE, where the PCA normalization is assumed in the reported fluxes.

When fitting Cen A data using *XSPEC* 12.5.1n (which allows for use of background

correction files for more than one detector), the correction factor for imperfect knowledge of the PCU2 background and PCU2 and HEXTE dead time estimates was an integral part of the fitting procedure, and in this manner uncertainties in the correction were included in the calculation of parameter uncertainties. The use of this procedure required that the XSPEC parameter `delta`, which defines the step size used in numerical determination of the derivatives used in the fitting process, be set to  $1 \times 10^{-4}$  (**K. Arnaud, private communication**) and that numerical differentiation be used (its value in the `.xspec/Xspec.init` file must be set to `true`<sup>3</sup>). The resulting negative values of the PCU2 corrections for the first 11 observation intervals ranged from a few to less than 10%, which is consistent with the values found by Rothschild et al. (2006) for the first 6 *RXTE*/PCU2 observations of Cen A. The 2009 January and February observations using the *SkyVLE* background required positive corrections of a few percent. Figure 2 shows the values of the background correction factor for the 13 observation intervals and the 47 daily data sets for both PCU2 and HEXTE. The corrections for the two different PCU2 background estimations are quite different.

## 4. Results

### 4.1. The Energy Spectrum of Cen A

All data were fitted with a model representing red-shifted line of sight absorption (XSPEC model `ZPHABS`) with abundances set to those of Wilms, Allen & McCray (2000) and utilizing the Verner et al. (1996) cross sections, a power law (XSPEC model `PEGPWRLW`) with a normalization equal to the unabsorbed flux in the 2–10 keV range, and a red-shifted Gaussian component representing iron  $K\alpha$  emission with width set to the *Suzaku* value of 30 eV (Markowitz et al. 2007). Instrumental line-like residuals near 4.5, 8.0, and 29 keV, when present, were fit as Gaussians as part of the model. Including one or more of these

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<sup>3</sup><http://heasarc.gsfc.nasa.gov/docs/xanadu/xspec/manual/XSmodelRecorn.html>

Gaussians did not significantly affect the best-fit parameters of the rest of the model, but did reduce  $\chi^2$  nearer to 1 in most cases.

The best-fit spectral parameters for the power law fit to the 13 observation intervals are given in Table 2 and the 2–10 keV and 20–100 keV fluxes are given in Table 3. In order for comparisons to spectra modeled with the XSPEC model POWERLAW, the best-fit flux at 1 keV is also given in Table 3 for additional fitting done with the POWERLAW model component replacing the PEGPWRLW description of the power law. We achieved good fits with both versions of this model in all cases.

A second model was tested with the power law replaced with the XSPEC model CUTOFFPL, which is a power law times an exponential to approximate any rollover of the spectrum due to Comptonization processes. This second model did not result in significantly lower  $\chi^2$  values in any case, and since it contained an added parameter (the cutoff energy), we used the results from the pure power law fits in the discussions that follow. Lower limits, and in 2 cases best-fit values, to the rollover energy are given in Table 3 for comparison to such fits by other authors. The 2 observations yielding best-fit values of the rollover energy have a common range of 382 keV to 693 keV. This is consistent with the Rivers, Markowitz, & Rothschild (2010) lower limit of 490 keV from fitting the sum of all the Cen A data. The observations with the strongest lower limits to the cutoff energy ( $\geq 2796$  keV and  $\geq 2499$  keV), however, occur at the highest flux values in 2009. This latter result is strong evidence for no cutoff being necessary, as measured over the 3–200 keV band (Fig. 3). Future instrumentation with good sensitivity to hundreds of keV acquired over a realistic amount of observing time will be required to derive a meaningful value for the rollover energy, if indeed that is the proper description of the high energy spectrum of Cen A.

A third model included the XSPEC model PEXRAV component of Compton reflection from cold/neutral material in a semi-infinite slab (i.e., from the accretion disk; Magdziarz & Zdziarski 1995) along with the observed power law. The folding energy was set to 1000 keV and the inclination angle was fixed at  $62^\circ.6$  (Burtscher et al. 2010) from the dust

emission in the mid-infrared. This model also did not improve the fit over a single power law, but did provide upper limits on the reflection component (Table 3). All upper limits were consistent with the Rivers, Markowitz, & Rothschild (2010) upper limit of 0.5% from a similar fit to the entire Cen A set of observations.

#### 4.1.1. Time Average Cen A Spectra

In order to set the stage for studying spectral variability of Cen A in the next section, we present the results of a separate study of the time averaged spectra of 23 bright active galaxies observed by *RXTE* over the duration of the *RXTE* mission (Rivers, Markowitz, & Rothschild 2010). In the case of Cen A, the time average spectrum included all of the observations through 2009 February. Rivers, Markowitz, & Rothschild (2010) found a mean Cen A power law index  $\Gamma=1.83\pm 0.01$ , a mean  $N_{\text{H}}=(1.69\pm 0.03)\times 10^{23}$  cm<sup>-2</sup>, a mean iron line energy of  $6.38\pm 0.09$  keV, and a mean iron line flux of  $(4.9\pm 0.7)\times 10^{-4}$  photons cm<sup>-2</sup> s<sup>-1</sup>. The upper limit to a Compton reflection component was  $R<0.5\%$ , and the lower limit to the high energy rollover was  $>490$  keV. As will be demonstrated below, these values are consistent with the mean values found by averaging the 13 observation intervals, i.e., a mean power law index of  $1.822\pm 0.004$ , a mean iron line energy of  $6.361\pm 0.014$  keV, and a mean iron line flux of  $(4.55\pm 0.14)\times 10^{-4}$  photons cm<sup>-2</sup> s<sup>-1</sup>.

## 4.2. Spectral Variability

### 4.2.1. Months to Years

Individual Cen A observations spanned intervals from less than 1 ks to a few hours, with separations between observations of days to 3 years. Consequently the spectral analysis could be performed on a range of temporal scales. As will be shown, the best fit parameters of the day-to-day observations did not show large variability and thus only the best-fit

parameters from the 13 observational intervals are given in Table 2. Figures 4 to 8 display the 13 sets of best fit spectral parameters for the observation intervals on the left and those from fitting the individual days on the right. Figure 4 shows the best-fit unabsorbed power law 2–10 keV flux, column density, and power law index; Fig. 5 shows the best-fit iron line energy, flux, and equivalent width, and Fig. 6 shows the 2–10 keV and 20–100 keV total fluxes from Cen A.

The inferred line of sight column density experienced nearly a factor of two in variability, with an interval of high density in the 2003 March to 2004 February data and lower density otherwise. The power law index remained essentially constant ( $\Gamma=1.822\pm0.004$ ) over 12.5 years, with a slight ( $\sim 2\%$ ) reduction at the time of the increased column density. The best-fit iron line energy values ( $6.361\pm0.014$  keV) were consistent with fluorescence of cold/neutral material at 6.4 keV, and the iron line flux appears to vary minimally over the observing decade with a mean value of  $4.55\pm0.14$  photons  $\text{cm}^{-2} \text{s}^{-1}$ . The equivalent width of the iron emission ranged from  $46_{-8}^{+10}$  to  $155_{-35}^{+40}$  eV, and the variation was as expected for a constant flux and varying power law continuum responsible for the fluorescing flux.

Figure 7 (Top) shows the values of the iron line flux plotted versus the power law normalization for the 13 observing intervals (Left) and the 47 daily measurements (Right). The average iron line flux is given by the horizontal line indicating no overall dependence of the iron line flux with the power law normalization. Similarly, the power law index showed no systematic variation with instantaneous power law flux (Fig. 7-Bottom). The very sparse sampling of the *RXTE* observations over 12.5 years prevented a meaningful conclusion vis-à-vis the relation between the power law emission and fluorescent iron line fluxes, i.e., any time delay between power law and iron line fluxes.

Figure 8 (Top) shows the variations in iron line flux versus column depth for the 13 observing intervals and for the daily observations, while Fig. 8 (Bottom) displays the equivalent width of the iron line versus column depth. Fig. 9 shows that the equivalent width of the iron line is inversely correlated with the power law flux, which is to be expected

if the iron line flux is essentially constant with varying power law flux. Figure 10 reveals that there is no temporal correlation of the iron line with either the 2–10 keV power law flux or the column density. **In addition, Figure 10 (Bottom) plots the measured column density versus the measured power law flux, clearly showing that no correlation exists between these two parameters also.** Hence, we find no correlations between any of the spectral parameters, other than the anti-correlation of equivalent width of the iron line and unabsorbed power law flux.

#### 4.2.2. Day to Day

We have combined data from individual ObsIds to produce 47 daily spectra from 1996 August 8 to 2009 February 21. As can be seen in Table 4, several sets of one or more contiguous days are present. The data from late 2009 January 15 and early January 16 have been combined into a single day since the data is effectively contiguous. Similarly the data for 2009 January 18, 19, and 20 have been combined, as well as 2009 January 25, 26, 27, 28. From these data sets we can derive indications of the variations from day to day exhibited by Cen A. Additionally, one can compare spectral parameters from observations a few days apart.

*Column Density  $N_{\text{H}}$* : Fig. 4 (Middle/Right panel) shows the best-fit values of  $N_{\text{H}}$  for each of the 47 daily spectra with the 13 observational intervals indicated. Significant variations in  $N_{\text{H}}$  are not seen on a day to day scale. The change over the 3 weeks between 2009 February 2 and 20 represents a 6% rise, while the interval from 2004 January 4 to February 13 had a 6% drop in column density over 40 days. On the 8 month timescale, a drop of 10% is seen from 2006 December 15 to 2007 August 18.

*Power Law Photon Index  $\Gamma$* : The value of the power law index was clearly below average in 2003 and early 2004 (Fig. 4 Bottom/Right panel). This decrement amounts to only 2% however.

*Power Law Flux:* The unabsorbed power law 2–10 keV flux is seen to vary with the observational interval fluxes (Fig. 4 Top/Right panel) in general. Within each interval, the daily values remain relatively constant, except for the January 2009 interval where the flux increased by about 20% with the power law index remaining constant. This variability is seen in both the 2–10 keV and 20–100 keV fluxes (Fig. 6), thus indicating that the power law intensity is the variable and not the power law index or absorbing column. **The lack of a correlation between the absorbing column density and the unabsorbed power law flux is clearly demonstrated in Fig. 10 (Bottom).**

*Iron Line:* The iron line centroid, flux, and equivalent width for each daily observation are shown in the right-hand panels of Fig. 5. The January 2004 values of the centroid are barely 1% above the average, while the flux is  $\sim 30\%$  above the mean. While this is simultaneous with the highest value of  $N_{\text{H}}$  and the 2% drop in the power law index, it is not considered significant when the overall variations are taken into account (see Fig. 10).

From this we can conclude, as with the 13 observational intervals, the Cen A spectrum only varies in power law normalization and column density, with relatively small variations in individual spectral parameters beyond that. Sustained monitoring at regular intervals over several months would be necessary for possible reverberation mapping of the iron line emitting material and for detailed studies of the variation in column density with respect to a clumpy torus model. Such monitoring by *RXTE* began in 2010 January and continues at present.

### 4.3. Temporal Variability

#### 4.3.1. Lightcurves

The PCU2 light curves of the first 6 observations (1996 August to 2004 February) were originally published in Rothschild et al. (2006), and the revised light curves for all observations are given here, reflecting improvements in the PCA response and background

estimation since then. Background subtracted light curves with 256 s time bins over the complete PCU2 energy band (2-60 keV) were generated for the 53 days with 4 or more time bins. In three cases observations overlapped slightly into the next day, and this is indicated in Table 4 (which gives the start and stop times of each daily light curve) by stop times greater than 24 hours. The resulting 51 light curves are plotted versus time since the time of the first temporal bin (Fig. 11). The short gaps in the light curves are the result of Earth occultations and passages through the SAA, while longer gaps and multiple day observations are the result of *RXTE* scheduling of other observations with higher priority. Also shown on the daily light curve figures are the average rate for the associated observation interval (dashed line) and the  $\pm 10\%$  rates (dot-dash lines).

It is clear from the figures that Cen A varies by  $\pm 10\%$  quite regularly over a week’s time, with some variations occurring on a daily timescale. This is especially clear in the 2004 January 2–4 and 2004 February 13–14 light curves. A 5% change over a few hours was seen between the end of 2006 December 12 and the beginning of December 13.

#### 4.3.2. Power Spectral Density Function

In this section, we present the broadband PSD analysis for Cen A by combining data from three X-ray instruments. While the resulting PSD does not have the same temporal frequency coverage as those for many previously-measured Seyferts, we were still able to derive a break in the power spectrum and demonstrate that the PSD shape was similar to those seen in both Seyferts and X-ray black hole binaries.

To facilitate PSD analysis, and allow use of the Discrete Fourier Transform (DFT) (Oppenheim & Schaffer 1975), we used only light curves which are evenly-spaced and relatively continuous (no large gaps). As shown below, for Cen A, we combined data from *RXTE*/All Sky Monitor (ASM), *RXTE*/PCA, and *XMM-Newton* to probe low, medium, and high temporal frequencies, respectively.



The sampling obtained by *RXTE/PCA* was relatively inhomogeneous, and we found only one “high-quality” PCA light curve which, after binning on the satellite orbital timescale, would make using analysis via a DFT straightforward: **that from the 2004 January 2–4 observation. These data were obtained from observation IDs 70152-01-01-14, 70152-01-02-000, 70152-01-02-010, and 70152-01-02-[00-16]; the resulting light curve, binned to the satellite orbital timescale of 5.7 ks, yielded a light curve with only 3/36 points missing.** With a duration of 2.2 days, we were able to probe variability on temporal frequencies near  $10^{-5} - 10^{-4}$  Hz.

*RXTE* is in a low-Earth orbit, and the resulting Earth occultation of the **PCA** every  $\sim 6$  ks severely complicates PSD analysis for temporal frequencies above  $1/(2 \times 6 \text{ ks})$ . For the high temporal frequency PSD ( $10^{-4} - 10^{-3}$  Hz), we relied on the uninterrupted light curves provided by *XMM-Newton* European Photon Imaging Camera (EPIC) pn. *XMM-Newton* observed the nucleus of Cen A twice, once in 2001 February for a duration of 23.4 ks (good exposure time 19.4 ks after screening), and again in 2002 February for a duration of 15.3 ks (good exposure time 8.9 ks). We used the longer-duration 2001 data (ObsID 0093650201; data obtained from the HEASARC public archive), which used the medium filter. Using XSELECT version 2.4a, we extracted 2–10 keV source and background light curves. **We used only PATTERN=0 events to reduce the impact of pile-up.** The source region was a circle of radius  $40''$  centered on the source; the background was extracted from an identical size circle  $\sim 3'$  away on the same CCD chip. We also extracted and inspected the 10–13 keV pn background light curve for flares, but found none. The light curves were binned to 300 s; variability at shorter time scales was dominated by Poisson noise.

The decade-long light curve from pointed PCA observations is too gap-dominated for reliable PSD analysis using a DFT, and so we rely on the light curve from the ASM. We downloaded 1-day averaged sum-band (1.5–12 keV) light curves from the MIT ASM database (<http://xte.mit.edu>), obtained between MJD 50087–54976. We removed those few

1-day data points which had negative fluxes or uncertainties greater than the average flux value, and binned the light curve to 20 days; power at temporal frequencies higher than  $\sim 1/(2 \times 20\text{d})$  was dominated by power due to Poisson noise and systematics associated with background subtraction, source confusion, etc. PSD measurements using 1- and 5-d binned light curves allowed us to empirically determine this level of power,  $P_{\text{Psn}}$ ; fitting a constant to the binned PSD above  $10^{-6.6}$  Hz yielded  $P_{\text{Psn}} = 33100 \text{ Hz}^{-1}$ .

Periodograms were measured for each light curve separately, binned into a PSD, and then the three PSD segments were combined (e.g., Edelson & Nandra 1999). Due to PSD measurement distortion effects (namely, aliasing, which affects only the PCA light curve, and red-noise leakage), the model-dependent Monte Carlo method described by Uttley, McHardy, & Papadakis (2002) **was used to assign proper uncertainties to each binned PSD point and determine the intrinsic, underlying PSD shape. The reader is referred to Uttley et al. (2002) for the definitions of the  $\chi_{\text{dist}}^2$  statistic used to compare observed and modeled PSDs and the rejection probability used to determine goodness of fit of model PSD shapes tested.**

Initial PSD construction closely followed §3.1 of Markowitz et al. (2003). Light curves were linearly interpolated across gaps, though such gaps were rare. Each light curve’s mean was subtracted. Following Papadakis & Lawrence (1993) and Vaughan (2005), the periodogram was logarithmically binned by a factor of  $\sim 0.20$  in the logarithm (roughly a **factor of 1.6** in  $f$ ) to produce the observed PSD,  $P(f)$ ; the two lowest temporal frequency bins were widened to accommodate three periodogram points. The constant level of power due to Poisson noise was not subtracted from these PSDs, but instead modeled in the Monte Carlo analysis. The individual long-, medium-, and short-term PSDs were combined to yield the final, broadband observed PSD, which is shown in Fig. 12(a). The PSD normalization of Miyamoto et al. (1991) and van der Klis (1997) was used to permit combining PSD segments from different missions; no additional renormalization of the individual measured PSDs was done.

We first tested an unbroken power-law model of the form  $P(f) = A_0(f/f_0)^{-\alpha}$ , where  $\alpha$  is the power-law slope and the normalization  $A_0$  is the PSD amplitude at  $f_0$ , arbitrarily chosen to be  $10^{-6}$  Hz. We stepped through  $\alpha$  from 0.0 to 3.2 in increments of 0.01. The best-fit model, plotted in Figure 12(a) as the dotted lines, was obtained for  $\alpha = 1.73 \pm 0.18$ . The rejection probability  $R_{\text{unbr}}$  was 0.948. The residuals are plotted in Figure 12(b). The errors reported here correspond to a value  $1\sigma$  above  $R_{\text{unbr}}$  for the best-fit value on a Gaussian probability distribution. For instance,  $R_{\text{unbr}}$  corresponds to  $1.94\sigma$ ; the error on  $\alpha$  corresponds to  $2.94\sigma$  or  $R_{\text{unbr}} = 0.997$ . (However, see Mueller & Madejski (2009) for warnings regarding using rejection probabilities as confidence regions for PSD model parameters.)

We then tested a singly-broken PSD model shape of the form  $P(f) = A_1(f/f_b)^{-\alpha_{\text{lo}}}$  for  $f \leq f_b$ , or  $A_1(f/f_b)^{-\alpha_{\text{hi}}}$  for  $f > f_b$ . where the normalization  $A_1$  is the PSD amplitude at the break frequency  $f_b$ , and  $\alpha_{\text{lo}}$  and  $\alpha_{\text{hi}}$  are the low- and high-frequency power law slopes, respectively, with the constraint  $\alpha_{\text{lo}} < \alpha_{\text{hi}}$ . PSD slopes were tested in increments of 0.1, as was  $\log(f_b)$ . The limited amount of PSD data precluded testing more complex PSD shapes.

The best-fit model had a rejection probability  $R_{\text{brkn}}$  of 0.208 for  $\log(f_b) = -(6.2^{+0.3}_{-0.2})$  (corresponding to  $T_b = 18.3^{+18.3}_{-6.7}$  d),  $\alpha_{\text{hi}} = 2.5^{+0.4}_{-0.1}$ , and  $\alpha_{\text{lo}} = 0.9^{+0.3}_{-0.2}$ . These errors correspond to values  $1\sigma$  above the rejection probability for the best-fit value on a Gaussian probability distribution; the best-fit model’s rejection probability corresponds to  $0.27\sigma$ ; the errors correspond to  $R_{\text{brkn}} = 0.796$  or  $1.27\sigma$  (Markowitz et al. 2003). The amplitude  $A_1$  was  $7.9^{+3.1}_{-2.3} \times 10^4$  Hz $^{-1}$ ; The best-fit value of  $A_1 \times f_b$  is 0.05, broadly consistent with values of  $\sim 0.01$ – $0.02$  measured for Seyferts.

The best-fit model is shown in Figure 12(a) as a solid line, and as shown in Figure 12(c), the residuals are much smaller compared to the unbroken power-law fit, particularly in the ASM PSD segment. Following Markowitz et al. (2003), using the ratio of likelihoods of acceptance  $L_{\text{brkn}}/L_{\text{unbr}}$  ( $L_{\text{brkn}} \equiv 1 - R_{\text{brkn}}$ ;  $L_{\text{unbr}} \equiv 1 - R_{\text{unbr}}$ ), the PSD break was significant at  $\sim 15\sigma$ .

Can the flattening in PSD slope seen in the ASM PSD segment be an artifact? For an energy spectrum with the shape of that for Cen A falling off rapidly below  $\sim 3$  keV, the **spectrum-weighted effective area** of the ASM peaks near roughly 4 keV, while that for the PCA peaks near 5–6 keV, so the difference in average photon energy may play a minor role at best. PSD power-law slopes above the break have been observed to flatten with increasing photon energy for an assumed energy-independent  $f_b$  (e.g., Nandra & Papadakis 2001; Vaughan & Fabian 2003), but the effect is in the opposite sense to that observed in the Cen A PSD. Finally, as the ASM light curve is virtually continuous, a constant level of power due to aliasing is not expected in this PSD segment. We conclude that the break in the PSD is real. Current *RXTE* monitoring to bridge the “gap” in the current PSD and cover the temporal frequency range  $10^{-6.6}$  to  $10^{-5.2}$  Hz is on-going and will be useful in further constraining the break frequency and PSD power-law slopes.

Finally, the effect of the observed evolution in line of sight column density  $N_H$  on the ASM PSD can be estimated as follows. The increase in  $N_H$  corresponds to a reduction in observed 2–10 keV flux by  $\sim 35\%$ . We added a  $\sim 2$ -year long trend with flux increasing/decreasing by 35% during 2003–2004 to the binned ASM light curve as a rough estimate of the intrinsic or unabsorbed light curve. The observed 2–10 keV flux of Cen A during 2003–2004 was below the long-term average, and so this action changed the fractional variability amplitude  $F_{\text{var}}$  (Vaughan et al. 2003) only slightly: a decrease of 3% from  $F_{\text{var}} = 29\%$  to 26%. That is, the unabsorbed light curve is slightly less variable than the observed light curve; the intrinsic PSD, compared to the observed PSD, may be lower by up to 17% in linear space, or 0.07 in log space, on temporal frequencies  $\lesssim 1.0 \times 10^{-8}$  Hz. In other words, the intrinsic broadband PSD may bend even more strongly than the observed PSD, albeit by a very small amount.

## 5. Discussion and Conclusions

A summary of the results from the *RXTE* observations includes: (1) the power law index was constant (to within a few percent) and independent of **2–10 keV power law flux, which varied by a factor of 3**, (2) the column density experienced a factor of 2 variation within a 1–4 year time span, (3) the iron line flux was independent of the power law flux, (4) the equivalent width was not proportional to the iron line flux, but varied inversely with the power law flux, (5) the iron line flux was not correlated with the column density variations, (6) a Compton reflection component was not detected, and (7) any high energy spectral roll over was greater than  $\sim 500$  keV and could be as high as 2 MeV.

In order to establish a picture of the geometry of the X-ray continuum emitter and the circumnuclear accreting material, these results must be combined with results from sub-parsec scale mid-infrared observations that reveal a nuclear, molecular torus with 0.6 pc diameter and a  $62^\circ.6$  inclination dusty disk (Burtscher et al. 2010), theoretical work describing the torus as a clumpy medium (Nenkova, Ivesi $\acute{c}$ , & Elitzur 2002; Nenkova et al 2008a; Nenkova et al. 2008b), and when combined with the steeper spectra seen at GeV and TeV energies, the combined X-ray/gamma ray data define the general shape and location of the blazar inverse Compton peak.

In the following sections we discuss the implications of the iron line diagnostics of the circumnuclear gas, inferences about a clumpy torus, the X-ray variability, the implications of not observing a correlation between the power law index and the power law flux, and the possible origin of the continuum emission. Our overall conclusions are then presented.

### 5.1. Fe $K\alpha$ Line Diagnostics of the Circumnuclear Gas

The lack of a strong observed Compton reflection component suggests that the bulk of the circumnuclear material is Compton-thin, or, if there does exist Compton-thick circumnuclear gas, it does not contribute significantly to the observed spectrum and/or is

poorly illuminated by the central X-ray source. In either case, the bulk of the observed Fe  $K\alpha$  emission line originates in Compton-thin gas, with at most a negligible contribution from Compton-thick gas.

The Fe emission line’s observed equivalent width  $EW_{\text{obs}}$  can be used as a diagnostic of the geometry of the Fe-line emitting gas under certain assumptions. For the moment, let us assume a simplified geometry in which the Fe-line emitting gas is distributed in an optically-thin, uniform spherical shell with line of sight column density  $N_{\text{H}}$  and a covering fraction as seen from the central X-ray continuum source  $f_{\text{cov}}$ . We use  $EW_{\text{obs}} = 90 \pm 10$  eV, the long-term average obtained by Rivers, Markowitz, & Rothschild (2010) from fits to the total summed PCA + HEXTE spectrum. We can then use Eq. 5 from Murphy & Yaqoob (2009) to relate  $EW_{\text{obs}}$  to  $N_{\text{H}}$  and  $f_{\text{cov}}$ . Assuming solar abundances ( $Z_{\text{Fe}} = 1$ ), the abundances of Wilms, Allen & McCray (2000), and an illuminating continuum with  $\Gamma = 1.83$ , we find  $EW_{\text{obs}} = 40 \text{ eV} \times (N_{\text{H}}/10^{23} \text{ cm}^{-2}) \times f_{\text{cov}}$ . A value of  $N_{\text{H}}$  of  $2.1 \times 10^{23} \text{ cm}^{-2}$  along with  $f_{\text{cov}} = 1$  will yield  $EW_{\text{obs}} = 90 \text{ eV}$ . Similarly, if the long-term average observed line of sight  $N_{\text{H}}$  from Rivers, Markowitz, & Rothschild (2010),  $1.5 \times 10^{23} \text{ cm}^{-2}$ , is used, a covering fraction of unity and  $Z_{\text{Fe}} = 1.5$  will yield the same value of  $EW_{\text{obs}}$ . That is, for all values of  $N_{\text{H}}$  less than a few times  $10^{23} \text{ cm}^{-2}$ , covering fractions near unity are required. This may indicate that the Fe line emitting gas is highly spatially extended as seen from the central source, consistent with the observed lack of strong variability in the intensity of the Fe line between 1996 and 2009, with response to the most rapid continuum variations getting smeared out.

Alternatively, if the Fe line emitting gas is not distributed uniformly, the observed equivalent width can still be attained with gas having a column density near  $10^{24} \text{ cm}^{-2}$  and  $f_{\text{cov}}$  near one-fifth; a torus-like structure lying mostly out of our line of sight would be consistent with this scenario. Following Fig. 8 of Murphy & Yaqoob (2009), assuming a torus-like structure with  $N_{\text{H}}$  of a few  $\times 10^{23} \text{ cm}^{-2}$  inclined so that it is not intersecting our line of sight,  $EW_{\text{obs}} \sim 90 \text{ eV}$  can be obtained if  $Z_{\text{Fe}} \sim 2$ . If this torus does intersect our line

of sight (consistent with full-covering absorption observed in hard X-ray spectra of Cen A), an  $EW_{\text{obs}}$  near 100 eV can be attained with  $Z_{\text{Fe}} \sim 1$ .

We also applied the new model “MYTorus” (Murphy & Yaqoob 2009) to our time-averaged data. This model is self-consistent and includes Compton reflection, Fe line emission and line of sight absorption due to a dusty circumnuclear torus. It has the benefit of being a physical model but also has the drawback of requiring certain simplifications and assumptions made about the geometry of the torus (namely that the half-opening angle of the torus is  $60^\circ$ , the torus cross section is circular, and the torus is uniform in density). Note that “MYTorus” also uses the abundances of Anders & Grevesse (1989) which yield lower values of  $N_{\text{H}}$  than those of Wilms, Allen & McCray (2000), which are used throughout the above analyses. However it gives us certain insights into possible configurations of the material surrounding Cen A.

The scenario tested by MYTorus was that of a torus that intercepts the line of sight, accounting for the observed absorption. In this case we obtained an inclination angle for the torus of  $> 70^\circ$  to the line of sight, and a column density of  $N_{\text{H}} = 1.11 \pm 0.03 \times 10^{23} \text{ cm}^{-2}$ . The torus was able to account for only half of the Fe line flux observed and did not include any significant contribution from the Compton reflection hump. An additional toroidal component of material with column density  $N_{\text{H}} = 1.3 \pm 0.5 \times 10^{23} \text{ cm}^{-2}$  outside of the direct line of sight to the nucleus can account for the additional Fe line flux.

## 5.2. Inferences about a Clumpy Torus

Temporal variations in column density are very common in Seyfert galaxies on various timescales (Risaliti, Elvis, & Nicastro 2002; Lamer, Uttley, & McHardy 2003). In Cen A, the line of sight X-ray column density measured in 2003 March, 2004 January, and 2004 February was  $(2.3, 2.6, 2.4) \times 10^{23} \text{ cm}^{-2}$ , respectively, whereas before and after it was about  $1.6 \times 10^{23} \text{ cm}^{-2}$ . For the following discussion, we assume that this represented the passage

through our line of sight of a single clump of matter in the circumnuclear torus representing a change in  $N_{\text{H}} \sim 1 \times 10^{23} \text{ cm}^{-2}$ .

The clumpy torus model of Nenkova (Nenkova, Ivesik, & Elitzur 2002; Nenkova et al 2008a; Nenkova et al. 2008b), parameterizes the number of clumps along the line of sight at an angle  $\Theta$  to the torus equatorial pole as  $N_{\text{LOS}}(90^\circ - \Theta)$ . The model also characterizes the number of clumps along the line of sight as an exponentially decreasing function of  $\Theta$ , i.e., the maximum number of clumps is seen in the equatorial plane of the torus, and the number decreases with angle from the plane. If one clump passing through the Cen A line of sight has a column density that is a large fraction of the total column density seen, as is the case here, one can infer that the line of sight only contains a few clumps. This, in turn may be interpreted as viewing Cen A at an angle to the torus equatorial plane, such that our line of sight passes through the outer region of the torus.

Furthermore, using the angle of the torus to the line of sight as  $62^\circ.6$  (Burtscher et al. 2010) and  $N_{\text{LOS}}(90^\circ - \Theta) = N_0 \exp(-(\Theta/\sigma)^2)$ , where  $N_0$  is the average number of clouds along the equatorial direction and  $\sigma$  characterizes angular distribution of clouds (Ramos Almeida et al. 2009), we find for  $\Theta = 62^\circ.6$ ,  $\sigma = 60^\circ$  (Ramos Almeida et al. 2009), and  $N_{\text{LOS}}(90^\circ - \Theta) = 3$ , that  $N_0 = 8$ . This is in accordance with the statement in Nenkova et al. (2008b) that “ $N_0$  is likely no larger than  $\sim 10$ – $15$  at most”, since a larger number would produce an infrared bump that is not seen. Therefore, viewing Cen A through the outer edges of the torus is consistent with the clumpy torus model.

If  $N_{\text{H}}$  for a representative clump is  $1 \times 10^{23} \text{ cm}^{-2}$ , then eight of them along the torus equatorial plane implies that the material in the equatorial plane may be close to being Compton thick ( $\sim 10^{24} \text{ cm}^{-2}$ ), but this value will drop as one goes farther from the plane of the torus, and the majority of the torus will be Compton-thin. Thus, the Compton thick aspects of the circumnuclear torus may be restricted to near the equatorial plane, and would explain the lack of a detected Compton reflection component. In addition, we have speculated previously that the high energy radiation from the nucleus could be



anisotropic with very little if any of it illuminating the Compton thick material (Rothschild et al. 2006). A combination of a minimum of Compton thick material and an anisotropic radiation pattern may provide a more complete explanation for the lack of a Compton hump in Cen A.

From Fig. 4 (Middle) we can estimate how long the clump took to pass through our line of sight. Assuming a sharp transition in  $N_{\text{H}}$  as the clump passed into and out of the line of sight, we estimate a duration  $T$  of a minimum of  $\sim 1$  year and a maximum of  $\sim 4$  years. Assuming a torus diameter of 0.6 pc (Meisenheimer et al. 2007) and a black hole mass  $M_{\text{BH}}$  of  $6 \times 10^7 M_{\odot}$  (Cappellari et al. 2009; Neumayer et al. 2010), the Keplerian velocity  $V_{\text{K}}$  of the clump can be used to estimate its linear size, if one knew the radial position of the clump within the torus. Using radial values of the IR torus of  $R_{\text{K}} \sim 0.1$  to  $\sim 0.3$  pc, we find the velocity of the cloud  $V_{\text{K}} = (GM_{\text{BH}}/R_{\text{K}})^{0.5}$  to be on the order of  $10^3 \text{ km s}^{-1}$  and has a linear dimension  $L_{\text{c}} = V_{\text{K}} \times T$  on the order of  $3\text{--}12 \times 10^{13} \text{ m}$ ,  $0.001\text{--}0.004 \text{ pc}$ , or about one to four light-days. Given the measured column density of the clump and the derived linear dimension, an inferred number density  $n_{\text{c}} = N_{\text{H}}/L_{\text{c}}$  is on the order of  $1\text{--}3 \times 10^7 \text{ cm}^{-3}$ . This yields an approximate cloud mass  $M_{\text{Cloud}} = \frac{4}{3}\pi(\frac{1}{2}L_{\text{c}})^3 m_{\text{p}} n_{\text{c}} \approx 1 \times 10^{30} \text{ g}$ , or about one Jupiter mass, **where  $m_{\text{p}}$  is the proton mass**. The linear dimension of the cloud is about a thousand times larger than those determined by Risaliti (2009) for clouds in the broad line region, while the density is three to four orders of magnitude smaller. The duration of the transiting event here is much longer than those seen in, e.g., NGC 1365 (Risaliti et al. 2009a,b) and NGC 3227 (Lamer, Uttley, & McHardy 2003), which lasted for days to a few months, and were consistent with absorbing clouds in the optical broad line region of these objects. The lower density inferred from the Cen A absorption event could be due to the clouds in Cen A being significantly further away from the supermassive black hole than those inferred for NGC 1365 or NGC 3227.

### 5.3. X-ray Variability of Cen A

We can compare the break inferred in the observed PSD of Cen A with those breaks measured in the PSDs of other Seyfert AGN by using the empirical relation between  $T_b$ ,  $M_{\text{BH}}$  and  $\dot{m}$  quantified by McHardy et al. (2006). This exercise relies on the assumption that accretion onto the supermassive black hole in Cen A proceeds via the same mechanism and with the same efficiency of radiation as the Seyferts used in the sample of McHardy et al. (2006). We assume a black hole mass of  $6 \times 10^7 M_\odot$ , the average of the estimates discussed in Section 1. The long-term, unabsorbed, 2–10 keV luminosity,  $L_{2-10}$ , (Rivers, Markowitz, & Rothschild 2010) is  $8.6 \times 10^{41} \text{ erg s}^{-1}$ . **The bolometric luminosity  $L_{\text{bol}}$ , is estimated to be  $8 \times 10^{42} \text{ erg s}^{-1}$  (yielding an inferred accretion rate relative to Eddington of 0.1%), assuming a conversion factor  $L_{\text{bol}}/L_{2-10} = 9$  (Marconi et al. 2004).** The best-fit empirical relation of McHardy et al. (2006),  $\log(T_b(\text{days})) = 2.10 \log(M_{\text{BH}}/10^6 M_\odot) - 0.98 \log(L_{\text{bol}}/10^{44} \text{ erg s}^{-1}) - 2.32$ , yields a predicted break time scale  $T_{b,\text{pred}}$  of 317 days, which corresponds to a break frequency of  $3.7 \times 10^{-8} \text{ Hz}$ . The value of  $T_{b,\text{pred}}$  is a factor of  $\sim 17_{-13}^{+36}$  higher than the observed PSD break time scale ( $T_b = 18.3_{-6.7}^{+18.3}$ ) measured above for Cen A.

The sample used by McHardy et al. (2006) was derived from mainly radio-quiet Seyferts accreting at  $\dot{m}$  of a few percent or more, and it is thus conceivable that radio-loud sources and/or sources accreting at very low values of  $\dot{m}$  do not strictly adhere to the relation derived by McHardy et al. (2006). However, the measured PSDs for the radio-loud AGN 3C 120 and the low-luminosity AGN NGC 4258 seem to be at least roughly consistent with the relation. 3C 120 has  $M_{\text{BH}} \sim 5 \times 10^7 M_\odot$  and  $\dot{m} \sim 0.31$  (Vasudevan & Fabian 2009), which yields  $T_{b,\text{pred}} = 1.1$  days. Marshall et al. (2009) and Chatterjee et al. (2009) have each measured the X-ray PSD and found best-fit values of  $T_b = 6.5$  days and 1.2 days, respectively, consistent with  $T_{b,\text{pred}}$ , considering the errors on  $M_{\text{BH}}$ ,  $\dot{m}$ , and  $T_b$ . NGC 4258's black hole mass is well-determined via mega-masers, at  $3.9 \times 10^7 M_\odot$  (Herrnstein et al. 1999), and  $\dot{m} = 10^{-4}$  (Lasota et al. 1996), so that  $T_{b,\text{pred}}$  is about 2000 days, consistent

with the lower limit of  $T_b = 4.5$  d measured by Markowitz & Uttley (2005). We speculate that different physical mechanisms may dominate the observed X-ray variability in Seyferts and in radio-loud and/or low- $\dot{m}$  AGN such as Cen A. That is, in such objects, the X-ray emission may be dominated by emission from a jet component versus that coming from a corona associated with the accretion disk (as in 3C 120 and the radio-quiet Seyferts).

#### 5.4. Lack of a Power Law Flux — Power Law Index Correlation

As presented in Fig. 7, no correlation was seen in the plot of power law index versus unabsorbed 2–10 keV power law flux, when the flux varied over a factor of three ( $3\text{--}10 \times 10^{-10}$  ergs  $\text{cm}^{-2}$   $\text{s}^{-1}$ ). This is in contrast to what is commonly reported for Seyfert 1s in the 2–10 keV band (e.g., Shih, Iwasawa, & Fabian 2002; Papadakis et al. 2002), where the phenomenon of increasing flux being positively correlated with increasing power law indices is virtually ubiquitous. All of the objects for which this correlation was seen are radio quiet Seyferts. While the radio quiet Seyferts have a positive correlation, a relativistically broadened Fe  $K\alpha$  emission component in about 2/3 of the cases (Nandra et al. 2007), and significant Compton reflection components, Cen A exhibits none of these, and has a powerful jet. Perhaps the anisotropic radiation pattern at the base of the Cen A jet only illuminates the accretion disk minimally, if at all, but does illuminate a portion of the clumpy torus. This could explain the lack of a measured Compton reflection component from the disk, if the disk is indeed Compton-thick. If the power law/flux correlation originates from processes associated with soft photon production and subsequent up-scattering of them by the hot electrons in the accretion disk corona (e.g., Haardt & Maraschi 1993), the anisotropic flux distribution mentioned above would minimize this process in Cen A. This lack of a power law index/power law flux correlation may also be explained by a lack of soft photons due to the presence of a radiatively-inefficient accretion flow, (e.g., Narayan, Mahadevan, & Quataert 1998; Quataert 2001), **and thus a lack of disk-produced soft photons, which are the basis for Comptonization models of**

hard X-ray emission in Seyferts.

### 5.5. Continuum Emission from the Jet in Cen A

Chiaberge, Capetti, & Celotti (2001) constructed the first spectral energy distribution (SED) for Cen A showing that the SED consisted of two components: a synchrotron component in the sub-millimeter to near infrared and a second higher energy synchrotron self-Compton (SSC) component in the X-ray to gamma ray range, where the latter was defined by the *CGRO* results (Kinzer et al. 1995; Steinle et al. 1998). Later, Lubiński (2009) presented the *INTEGRAL*/IBIS spectrum of Cen A extending to 400 keV or greater as a single power law. The lack of a significant deviation from a power law up to energies as high as 2 MeV, as inferred from *RXTE*, can be interpreted in the misoriented BL Lac scenario (Chiaberge, Capetti, & Celotti 2001) as having the X-ray emission form the low energy side of the SSC hump, viewed at a large angle from the jet direction. Despite the overall SED being reminiscent of blazars, Cen A’s X-ray spectrum exhibits two signatures of Seyferts and quasars: a high column density of absorbing gas in the line of sight, and an iron emission line. In the Cen A case the bright hard X-ray continuum emission may originate at the base of the jet, yet from behind the absorbing line of sight material in contrast to what is commonly observed from blazars. A jet origin to the continuum has also been put forth for FR I radio galaxy cores by Evans et al. (2006), but **these authors** associate low absorption ( $<5 \times 10^{22} \text{ cm}^{-2}$ ) with sources having the hard X-ray emission associated with a jet. In this regard, Cen A may be a counter example.

The nature of the accretion flow near the black hole may contribute to defining the difference in nature between low-luminosity, radio loud AGN, such as Cen A, and radio-quiet Seyfert AGN, since a radiatively inefficient flow may be responsible for launching jet outflows (e. g., Körding, Fender, & Migliari 2006). In this case one would not have the soft photons for subsequent upscattering by the hot electrons in the accretion disk corona (Haardt & Maraschi 1993). In Cen A, the accretion disk may only supply matter to the

base of the jet and into the black hole, but may not make any contribution in terms of observational X-ray signatures.

## 5.6. Conclusions

Over more than a decade of observing, Cen A has shown remarkably stable spectral behavior. Both the power law index and the flux of the iron line have remained essentially constant, in contrast to the range in power law flux and column density, which have varied by a factor of 3 and 2 respectively. The stability of the power law index over the past 40 years (with the exception of the early 1970's when it flattened considerably Baity et al. (1981)), implies a stability **of the processes generating** the continuum at the base of the jet. The spectral aspects of these processes, i.e. the shape of the electron distribution and electron optical depth, must therefore be relatively insensitive to the variations in the accretion rate experienced at the base of the jet, which is assumed to be directly related to the overall flux observed. From this we infer that the mass accretion rate in the inner portions of the accretion disk ( $\dot{m} \approx L_X/L_{\text{Edd}} \sim 10^{-4}$ ) has remained within the range where the **radiation efficiency is** relatively unchanged even though the luminosity (a tracer of the mass accretion rate) has varied by a factor of 3.

Our observations with *RXTE* have led to the conclusions that we are viewing the supermassive black hole at the nucleus of this active radio galaxy through the outer portions of a clumpy molecular torus which is mainly Compton-thin, but may approach being Compton-thick only at or near its equator. The Compton thin torus **may then** occupy a large solid angle as viewed from the Cen A nucleus in order to be the site of the iron line emitting material.

The variability characteristics, when compared to Seyferts, imply that the accretion disk is not the location of the  $>3$  keV X-ray flux, but that synchrotron self-Compton **scattering** in the base of the jet may be responsible. In the case of Cen A, the bright

hard X-ray continuum emission may originate at the base of the jet, yet from behind the absorbing line of sight material in contrast to what is commonly observed from blazars.

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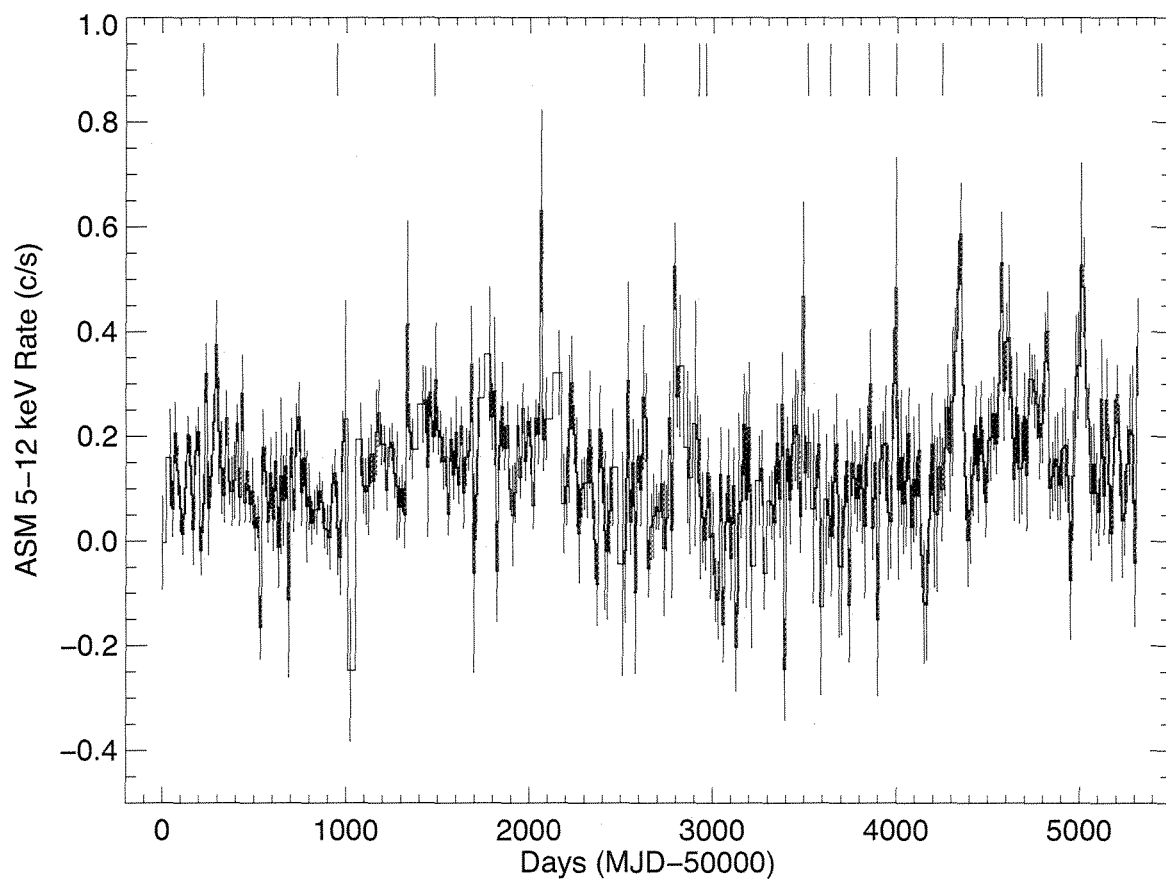


Fig. 1.— The flux history of Cen A over the *RXTE* mission as measured by the All-Sky Monitor 5–12.1 keV channel 3, with 2 week binning. The verticle bars at the top of the plot denote when the *RXTE* pointed observations were made.

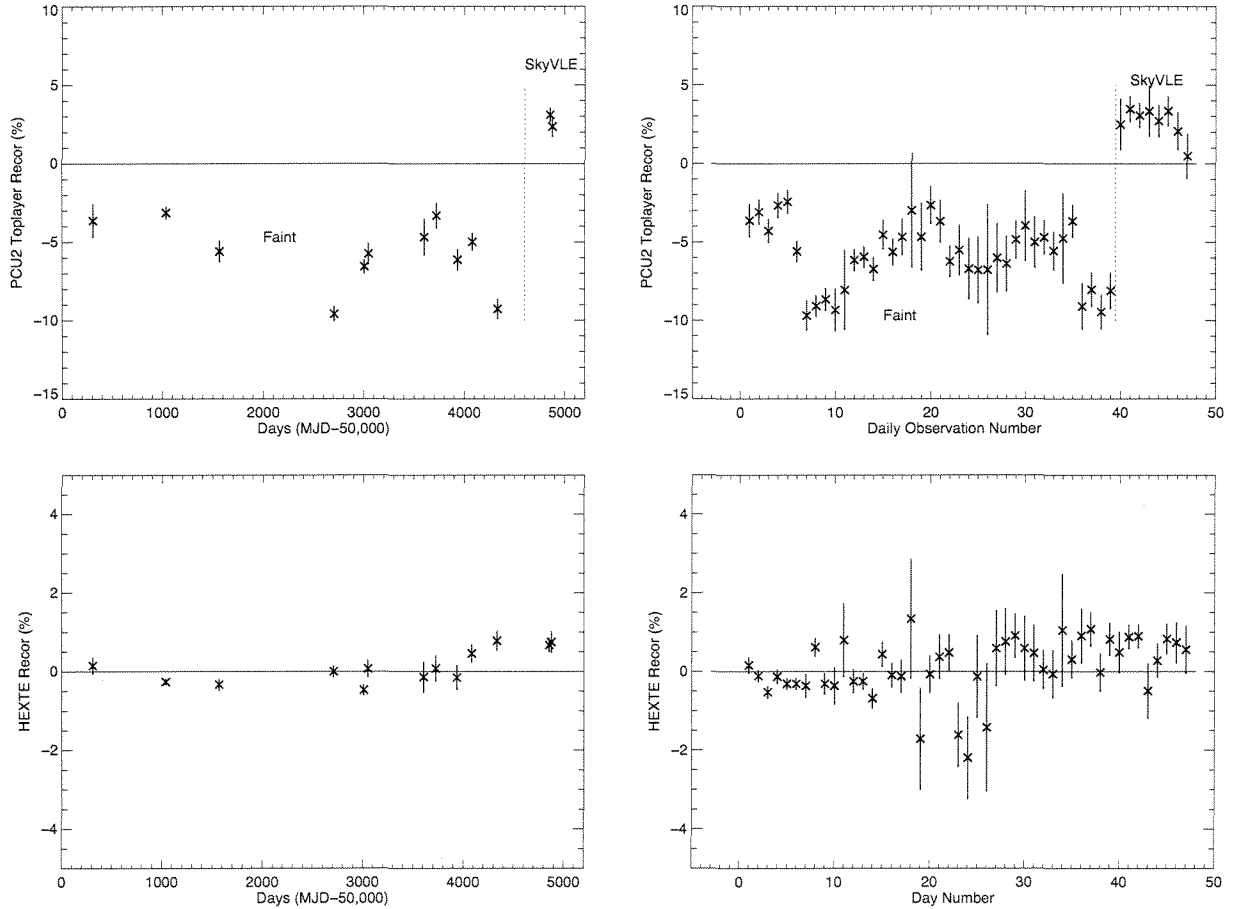


Fig. 2.— The best-fit correction factors applied to background/deadtime estimates for the 13 observation intervals and the 47 daily spectra histograms for PCU2 (Top) and HEXTE Bottom derived from fitting of Cen A spectra. The vertical dotted line in the PCU2 figures delineates the separation of Faint and SkyVLE background estimate use. Horizontal lines at a value of zero are included.

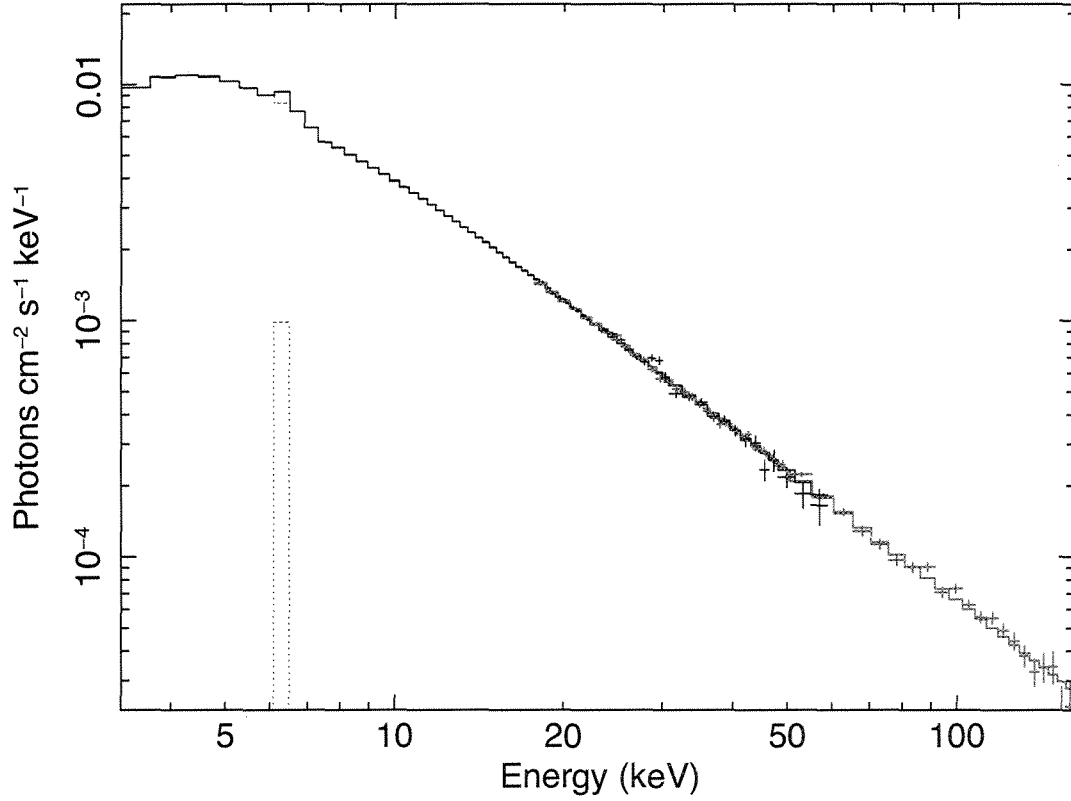


Fig. 3.— The inferred incident spectrum of the 2009 January set of observations. PCU2 toplayer 3–60 keV and HEXTE 18–200 keV data points are shown along with the model of a single absorbed power law with an iron emission line at 6.4 keV. The line component is indicated by the dotted lines.

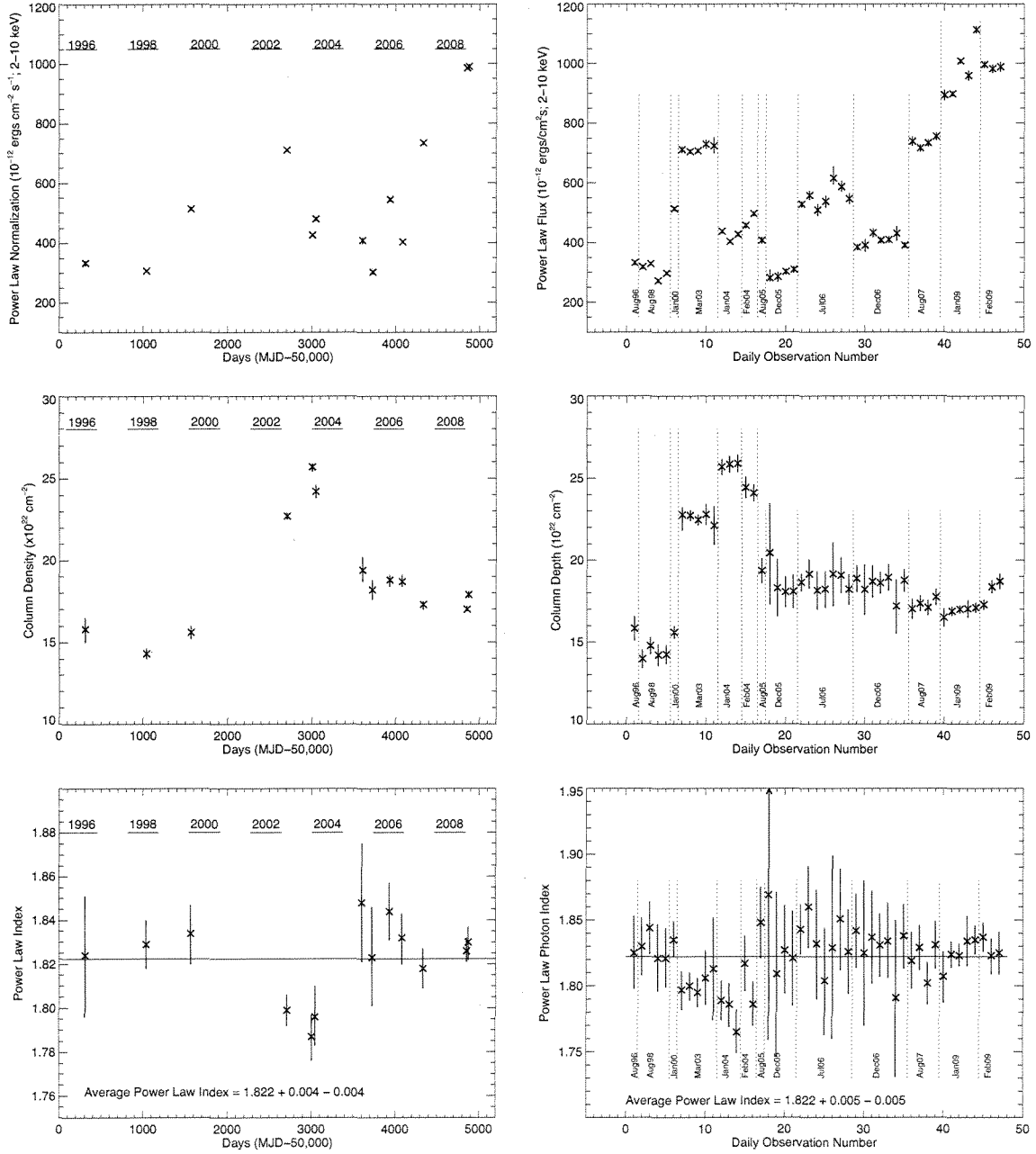


Fig. 4.— The best-fit 2–10 keV normalizations of the power law (Top), low energy absorption (Middle), and power law photon indices (Bottom) for the 13 observation intervals (Left) and the 47 daily spectra (Right). Years and observing intervals are shown. Average values plus 90% uncertainties are given for the power law index. Error bars represent 90% uncertainties.



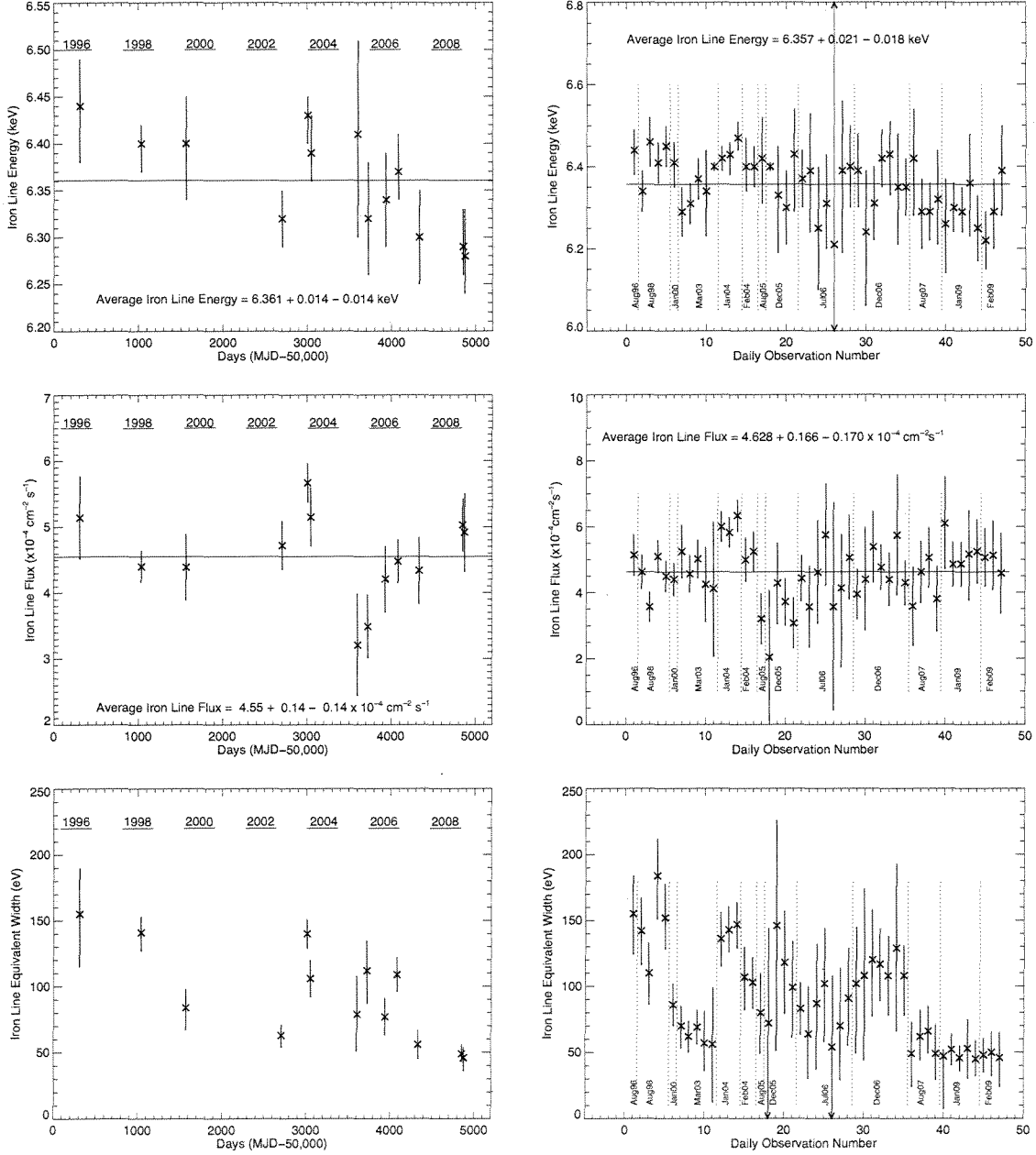


Fig. 5.— The best-fit iron line energy (Top), iron line flux (Middle), and equivalent width (Bottom) spectral parameters resulting from fitting the 13 separate observation interval spectra and the 47 daily spectra. Error bars represent 90% uncertainties. Average values plus 90% uncertainties are given for the iron line energy and flux. Years and observing intervals are shown.

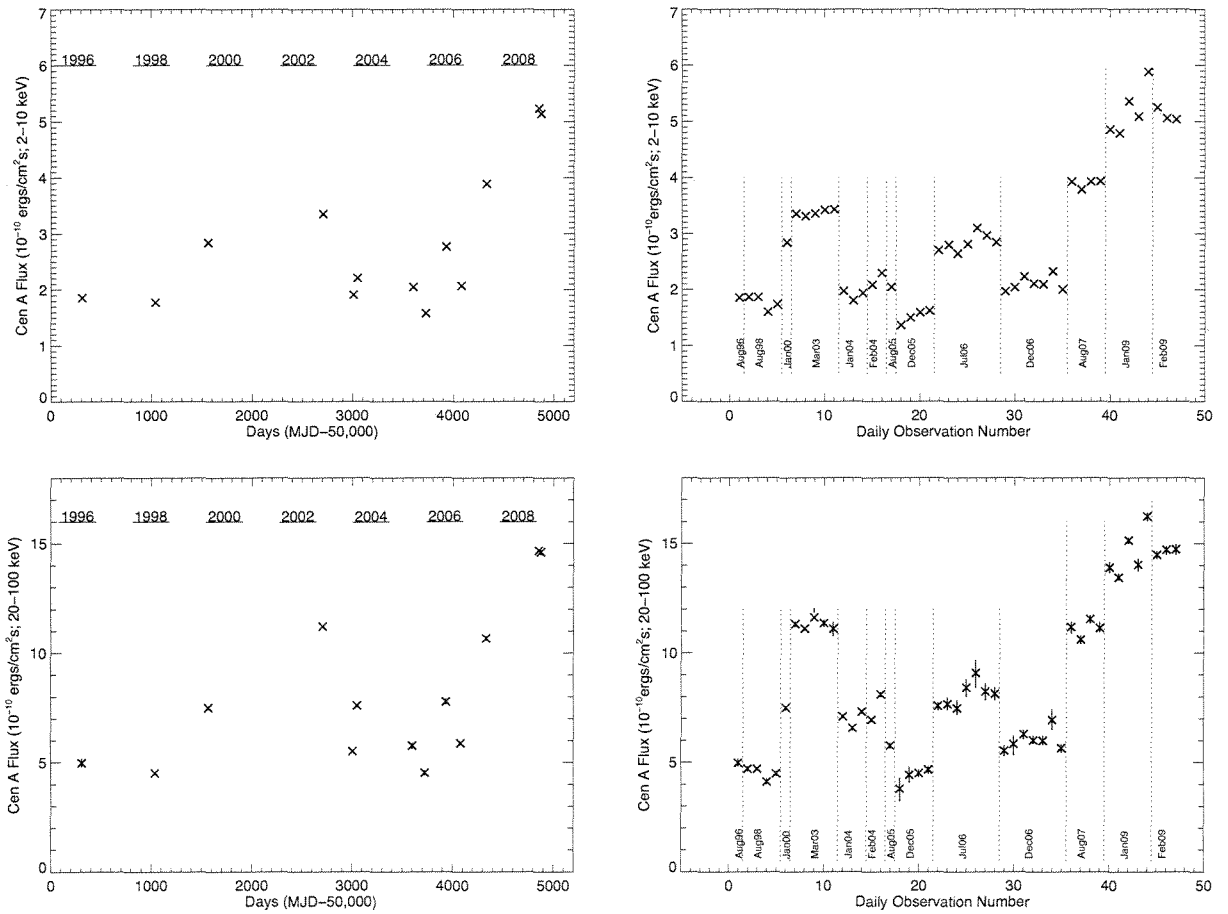


Fig. 6.— The *RXTE* Cen A 2–10 keV (Top) and 20–100 keV (Bottom) fluxes for the 13 observation intervals and the 47 daily spectra. Years and observing intervals are shown. Error bars representing 90% uncertainties are generally smaller than the symbols.

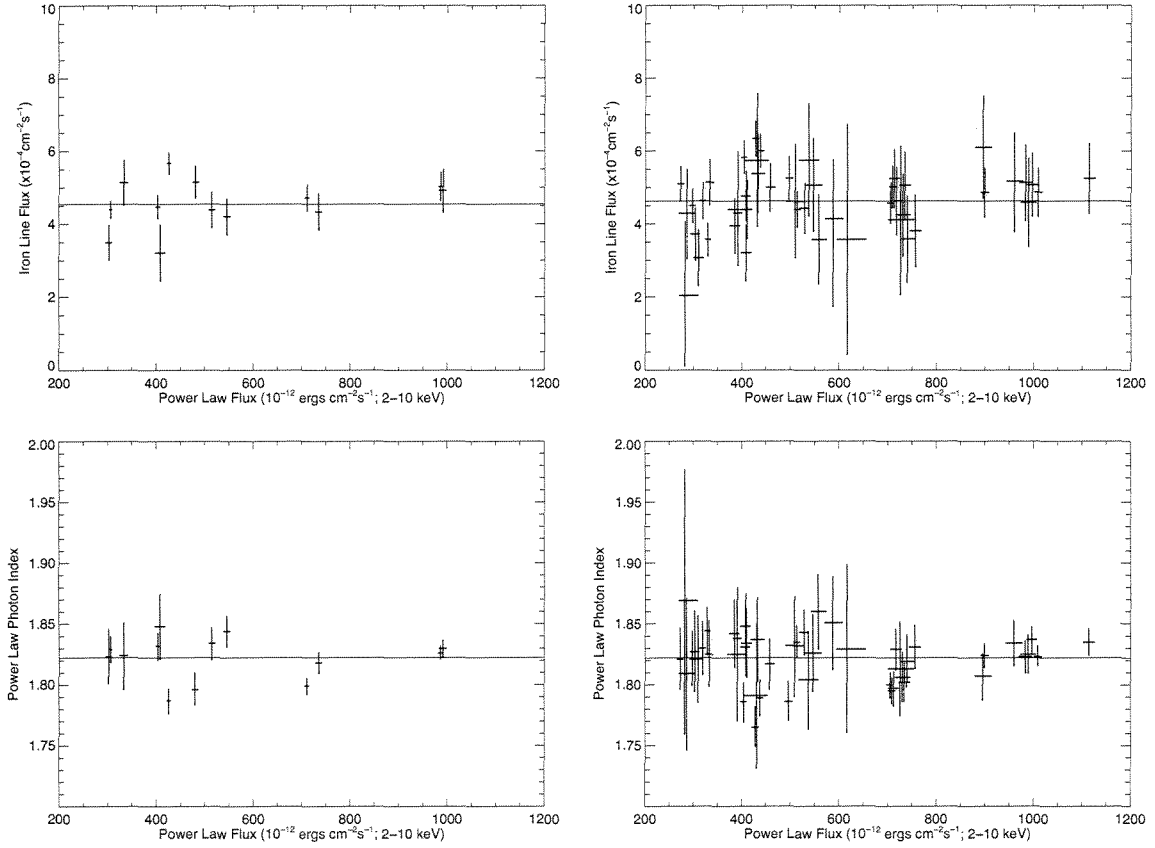


Fig. 7.— The iron line flux (Top) and power law index (Bottom) as a function of unabsorbed power law flux for the 13 observation intervals (Left) and 45 daily spectra (Right). The average iron line flux (Top) and average power law index (Bottom) are shown as a horizontal lines. Error bars represent 90% uncertainties.

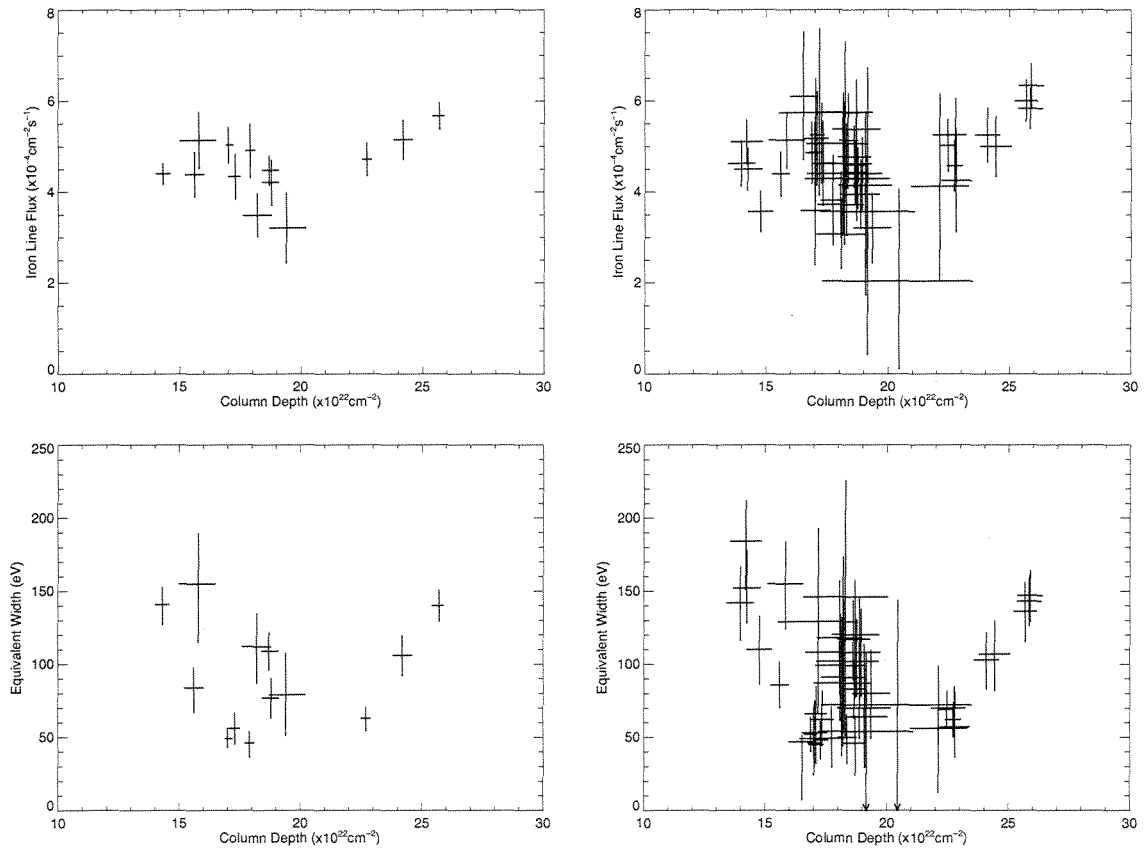


Fig. 8.— The iron line flux as a function of the low energy column depth (Top) and the iron line equivalent width as a function of the column depth (Bottom) for the 13 observation intervals (Left) and 47 daily spectra (Right). Error bars represent 90% uncertainties.

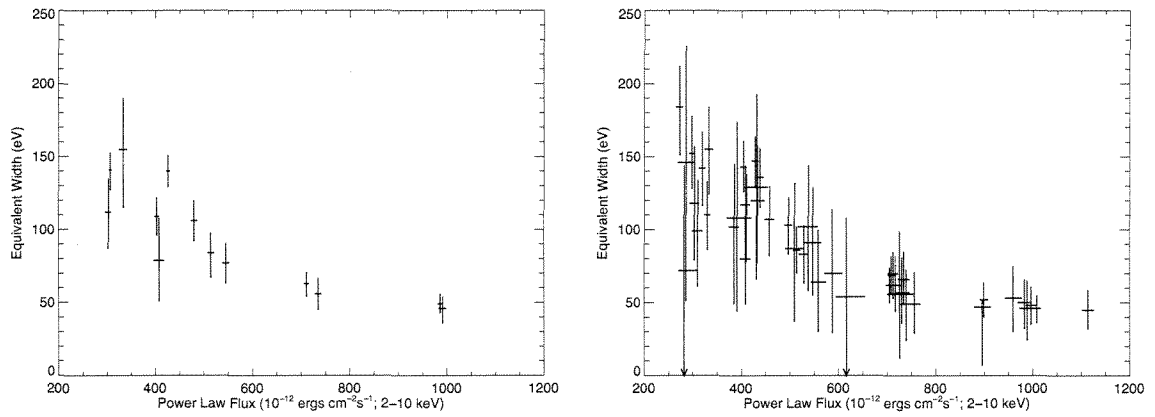


Fig. 9.— The equivalent width of the iron line is presented as a function of the unabsorbed power law 2-10 keV flux. The anti-correlation between the two parameters is clear, which indicates that the iron line was not responding to the power law continuum on the timescales available.

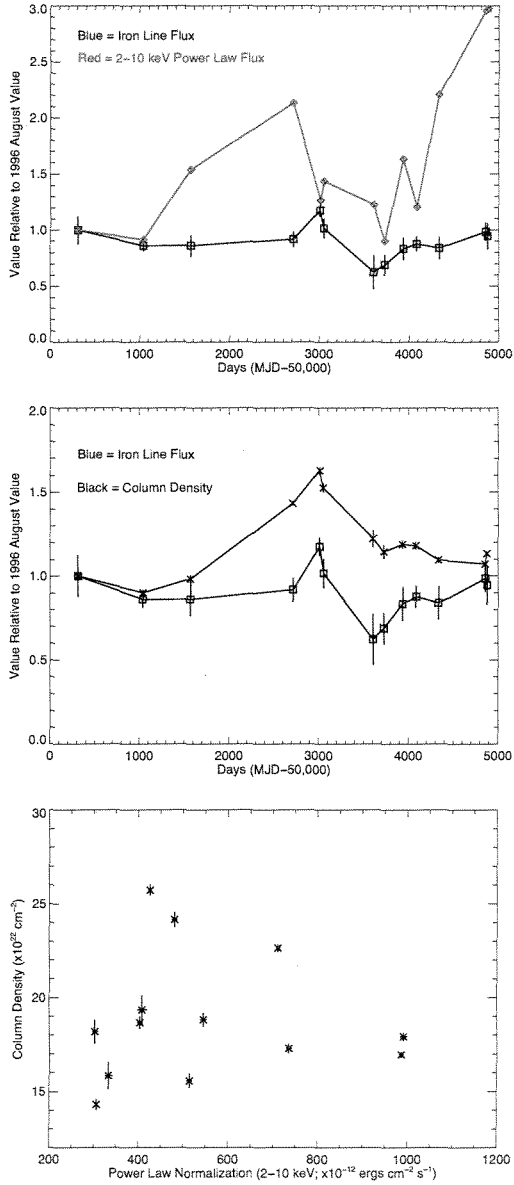


Fig. 10.— The variation in the iron line flux and the 2–10 keV unabsorbed power law flux (Top) and the column density and iron line flux (Middle) as a function of time. The values are normalized to the first observation in 1996 August. The lack of correlation is clear in both plots. The column density as a function of the power law flux (Bottom) shows no correlation between the two parameters.

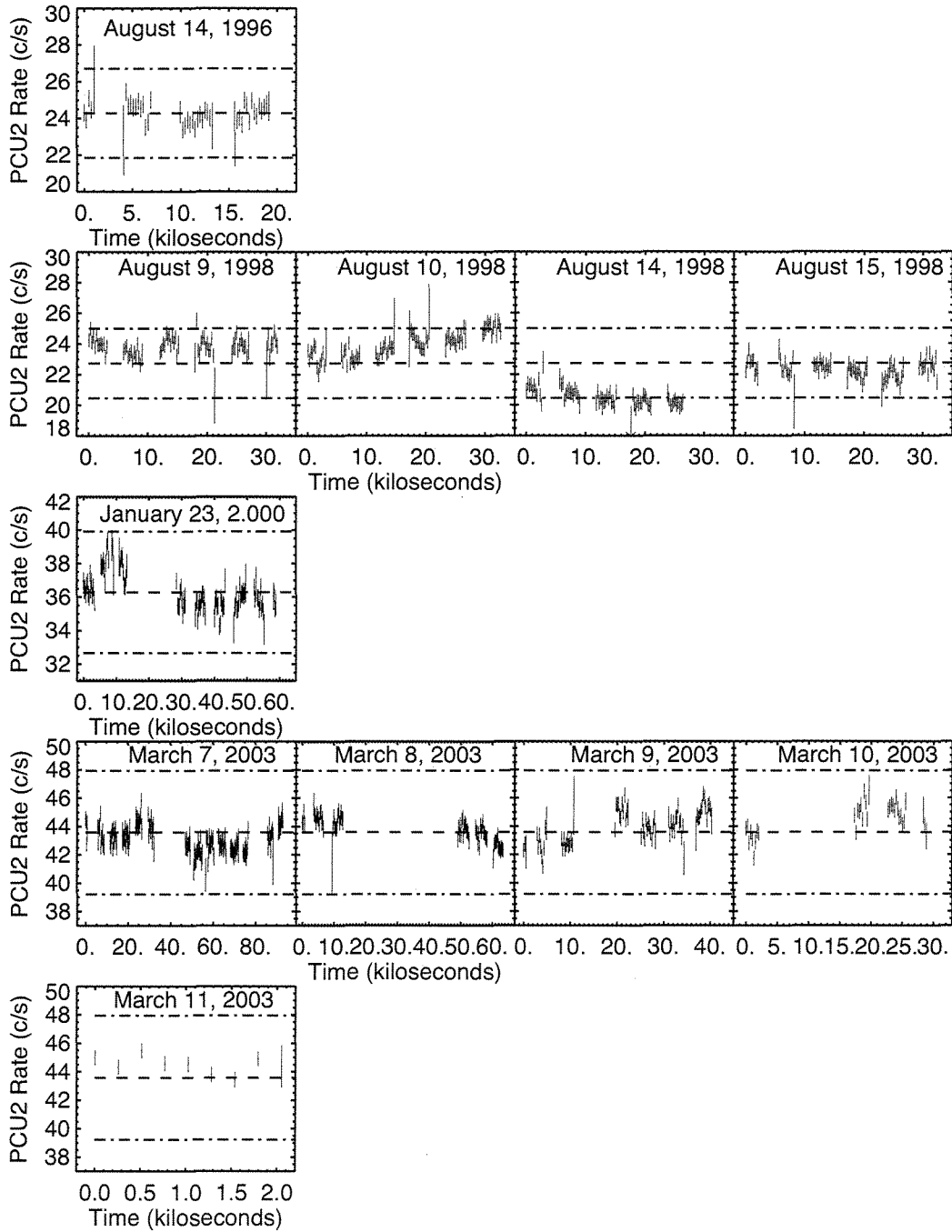


Fig. 11.— Daily light curves of Cen A with 256 s time bins over the full PCU2 energy range for observation intervals August 1996, August 1998, January 2000, and March 2003. Dashed line is the average rate for the observational interval and the dash-dot lines represent  $\pm 10\%$  in the rate.

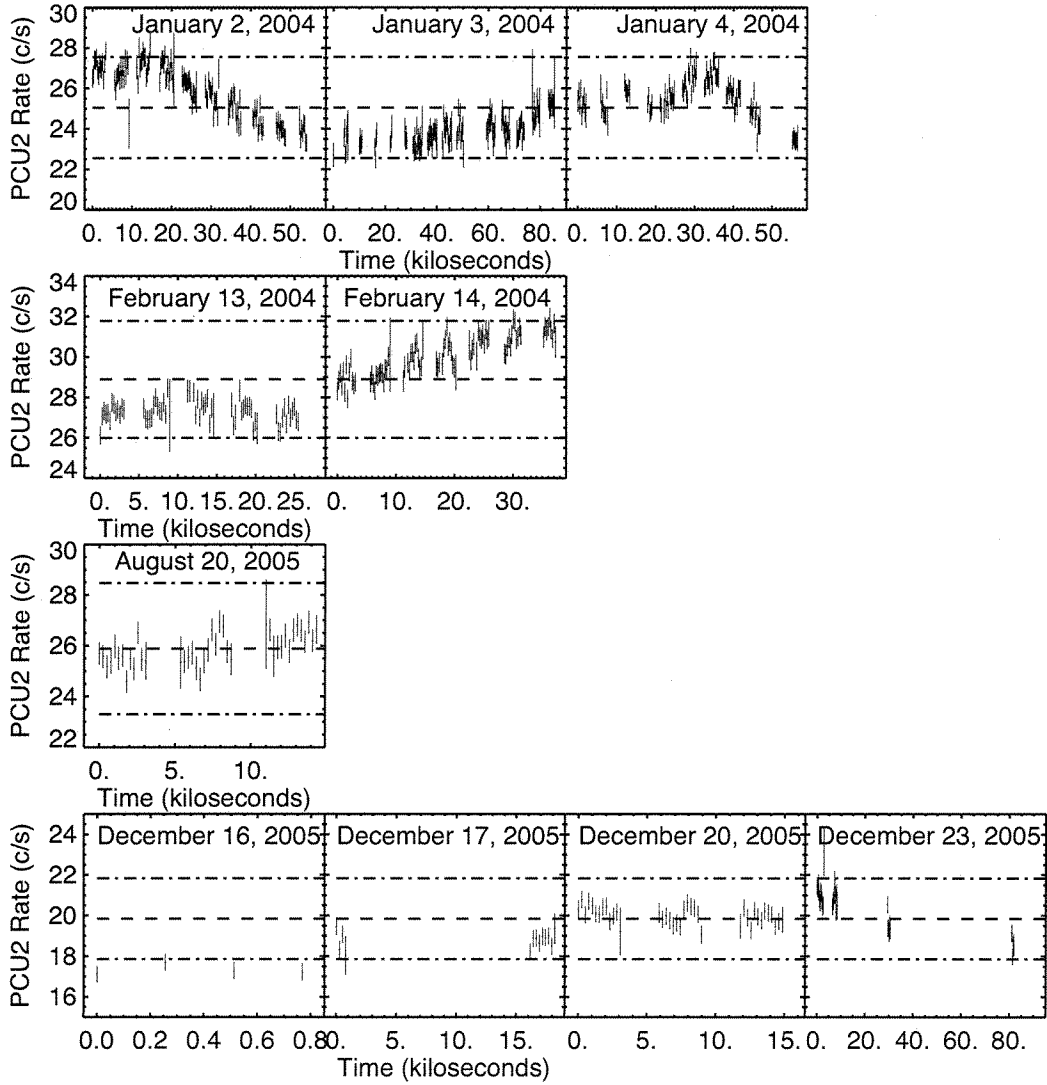


Fig. 11.— Daily light curves of Cen A with 256 s time bins over the full PCU2 energy range for observation intervals January 2004, February 2004, August 2005, and December 2005. Dashed line is the average rate for the observational interval and the dash-dot lines represent  $\pm 10\%$  in the rate.



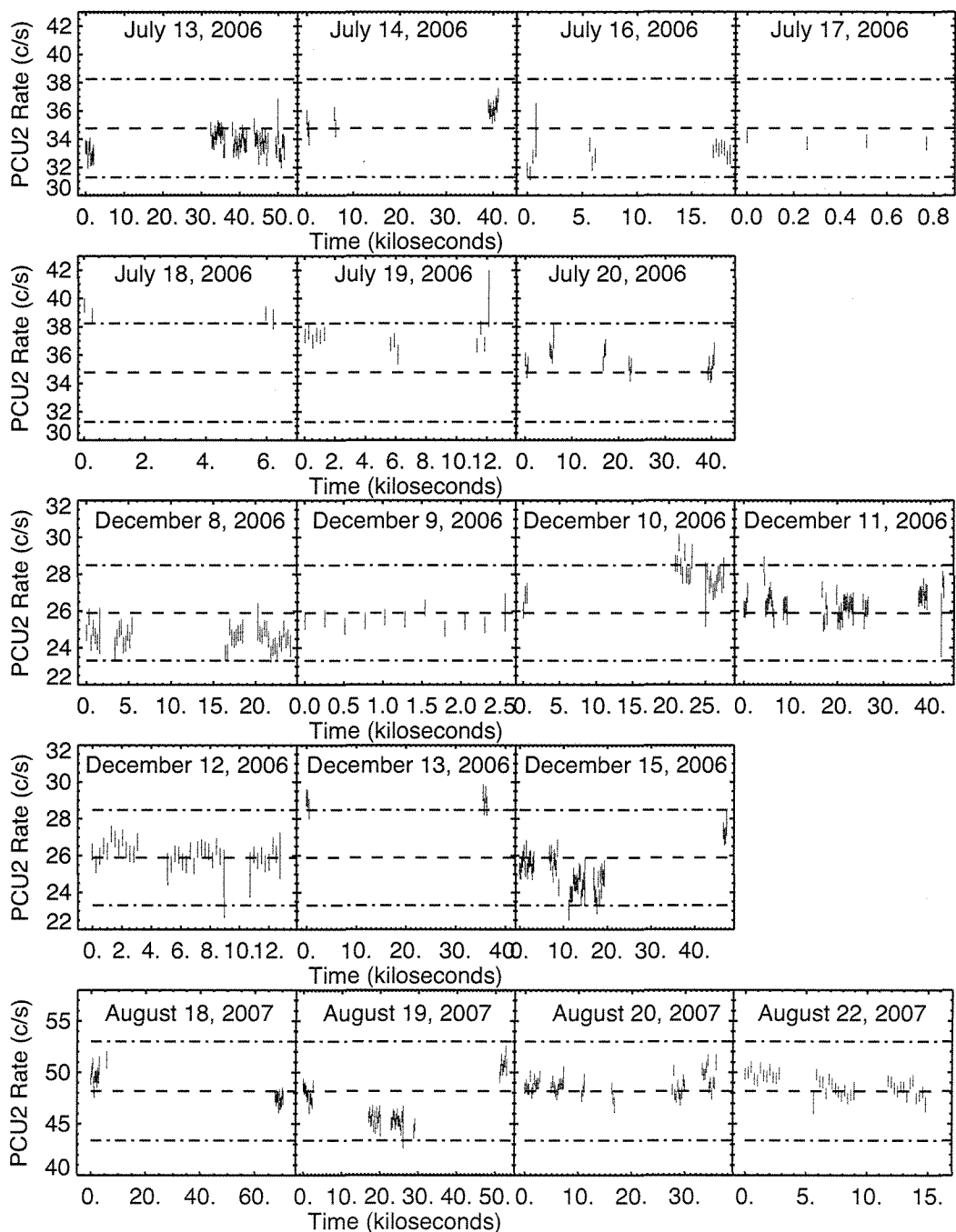


Fig. 11.— Daily light curves of Cen A with 256 s time bins over the full PCU2 energy range for observation intervals July 2006, December 2006, and August 2007. Dashed line is the average rate for the observational interval and the dash-dot lines represent  $\pm 10\%$  in the rate.

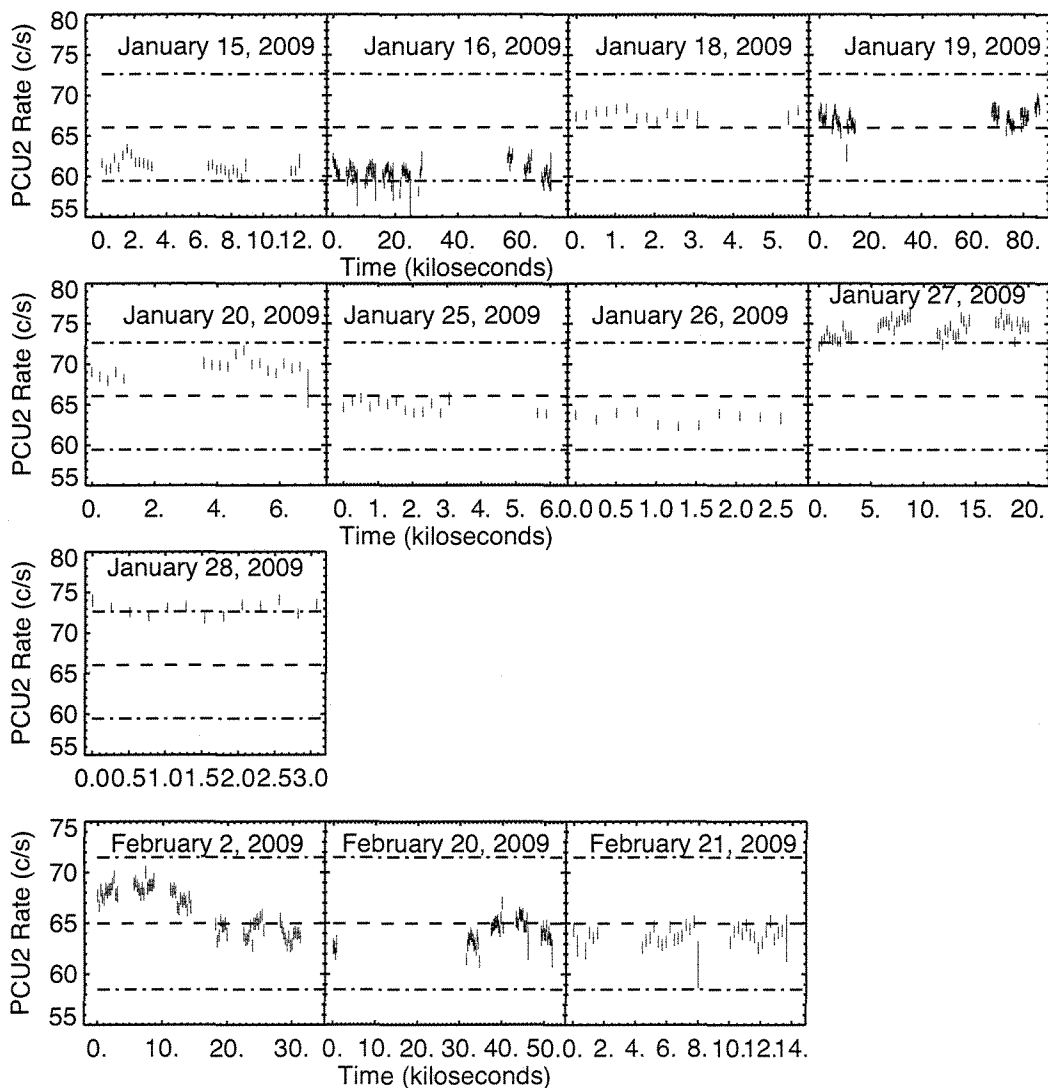


Fig. 11.— Daily light curves of Cen A with 256 s time bins over the full PCU2 energy range for observation intervals January 2009 and February 2009. Dashed line is the average rate for the observational interval and the dash-dot lines represent  $\pm 10\%$  in the rate.

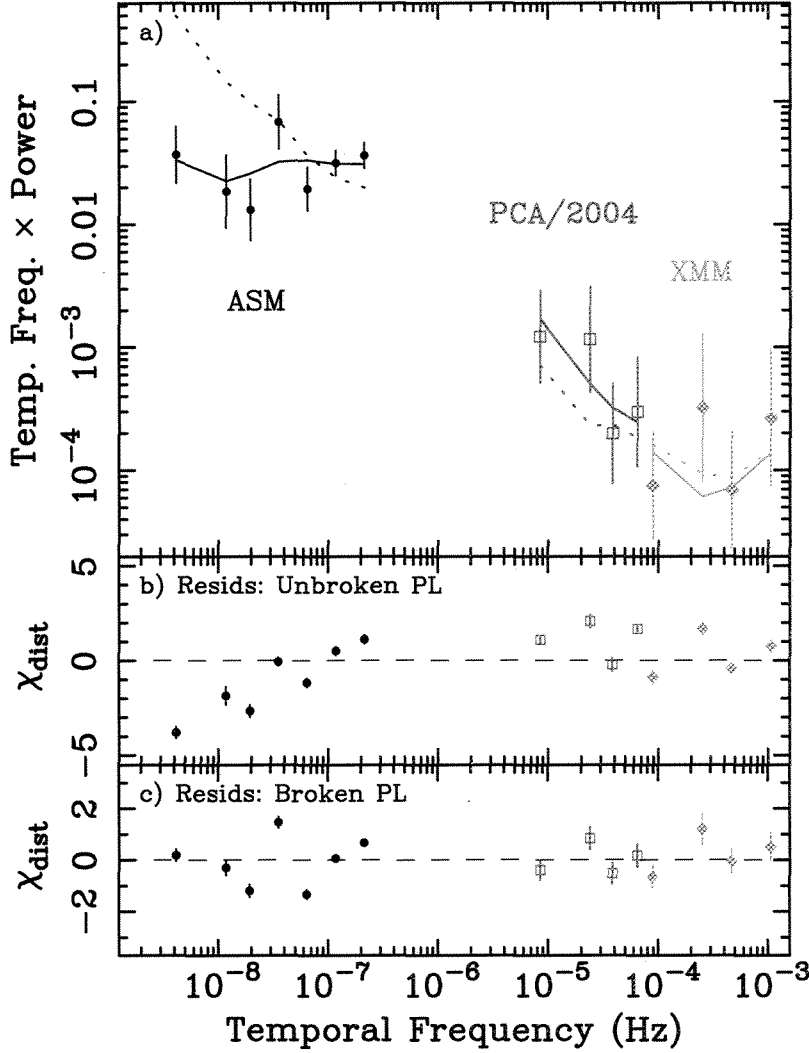


Fig. 12.— The power spectral density function for Cen A, derived from *RXTE* and *XMM* observations. In the top panel, the data points denote the observed binned PSD, with uncertainties on each point derived from the Monte Carlo simulation procedure (?). The dotted and solid lines denote the best fit unbroken and broken power law models, respectively; both are plotted in “data” space and contain the distortion effects inherent in PSD measurement. The  $\chi_{\text{dist}}$  residuals to the unbroken and singly broken power law fits are given in b) and c), respectively.  $\chi_{\text{dist}}$  compares the model PSD (from the Monte Carlo simulation procedure) and observed PSD, using the distribution estimated from the simulations (see Uttley, McHardy, & Papadakis 2002, for details).

Table 1. Livetimes and Rates for Observations of Cen A with *RXTE*

Date	PCU2 Livetime <sup>a</sup>	PCU2 Rate <sup>b,c</sup>	HEXTE-AB Livetime <sup>a</sup>	HEXTE-AB Rate <sup>b,c</sup>	HEXTE-B Livetime <sup>a</sup>	HEXTE-B Rate <sup>b,c</sup>	Num. Obs.	Num. Days
Aug96	10,528	23.24±0.07	6,785	5.12±0.20	3,390	4.88±0.27	1	1
Aug98	67,872	21.23±0.03	42,567	4.41±0.08	18,521	3.77±0.12	7	4
Jan00	25,088	34.58±0.05	16,132	7.72±0.12	8,037	7.01±0.15	3	1
Mar03	75,280	42.46±0.03	58,574	11.52±0.04	29,302	10.40±0.06	13	4
Jan04	88,960	24.46±0.04	59,788	6.86±0.04	26,015	6.13±0.05	22	3
Feb04	36,112	28.14±0.04	23,680	7.94±0.06	14,405	7.11±0.08	4	2
Aug05	9,472	24.95±0.07	6,073	6.47±0.13	3,021	5.53±0.17	1	1
Dec05	20,994	19.15±0.07	—	—	7,115	4.50±0.12	8	3
Jul06	29,968	33.46±0.04	—	—	9,228	7.04±0.10	22	7
Dec06	53,088	25.01±0.03	—	—	16,984	5.04±0.08	15	7
Aug07	38,784	46.50±0.04	—	—	12,997	10.44±0.10	11	4
Jan09	81,230	64.31±0.03	—	—	27,406	14.91±0.07	15	5
Feb09	38,800	63.90±0.05	—	—	12,171	15.03±0.10	7	3

<sup>a</sup>The Livetime in seconds

<sup>b</sup>PCU2 3–60 keV counting rate in counts/s; HEXTE 15–200 keV counting rate in counts/s

<sup>c</sup>Uncertainties in rates are 68% confidence

Table 2. Best-fit Parameters for the 13 Observations of Cen A with *RXTE*

	$N_H^a$	$\Gamma^b$	Norm(2-10) <sup>c</sup>	E(Fe K- $\alpha$ ) <sup>d</sup>	Flux(Fe K- $\alpha$ ) <sup>e</sup>	EW(Fe K- $\alpha$ ) <sup>f</sup>	$\chi^2/\text{dof}^g$	$\chi^2_\nu$
Aug96	$15.8^{+0.8}_{-0.7}$	$1.824^{+0.028}_{-0.027}$	$333.1^{+8.7}_{-8.3}$	$6.44^{+0.06}_{-0.05}$	$5.14^{+0.63}_{-0.63}$	$155^{+40}_{-35}$	135.8/118	1.15
Aug98	$14.3^{+0.3}_{-0.3}$	$1.829^{+0.011}_{-0.011}$	$305.4^{+3.1}_{-3.1}$	$6.40^{+0.03}_{-0.02}$	$4.40^{+0.24}_{-0.24}$	$141^{+14}_{-12}$	141.2/119	1.19
Jan00	$15.6^{+0.4}_{-0.4}$	$1.834^{+0.014}_{-0.013}$	$513.1^{+6.8}_{-6.6}$	$6.40^{+0.05}_{-0.06}$	$4.39^{+0.50}_{-0.50}$	$84^{+17}_{-14}$	118.3/118	1.00
Mar03 <sup>h</sup>	$22.7^{+0.2}_{-0.2}$	$1.799^{+0.007}_{-0.007}$	$711.8^{+5.0}_{-4.2}$	$6.32^{+0.03}_{-0.03}$	$4.72^{+0.37}_{-0.37}$	$63^{+9}_{-8}$	118.3/117	1.01
Jan04 <sup>i</sup>	$25.7^{+0.3}_{-0.3}$	$1.787^{+0.010}_{-0.010}$	$422.5^{+3.9}_{-3.9}$	$6.43^{+0.03}_{-0.02}$	$5.67^{+0.30}_{-0.30}$	$140^{+11}_{-11}$	144.9/116	1.25
Feb04 <sup>h</sup>	$24.2^{+0.4}_{-0.4}$	$1.796^{+0.013}_{-0.014}$	$478.6^{+5.7}_{-6.3}$	$6.39^{+0.03}_{-0.04}$	$5.15^{+0.44}_{-0.45}$	$106^{+14}_{-14}$	125.2/117	1.07
Aug05	$19.4^{+0.8}_{-0.8}$	$1.848^{+0.027}_{-0.027}$	$408.5^{+10.9}_{-10.5}$	$6.41^{+0.11}_{-0.10}$	$3.17^{+0.77}_{-0.77}$	$79^{+28}_{-29}$	96.4/118	0.82
Dec05	$18.2^{+0.6}_{-0.6}$	$1.823^{+0.022}_{-0.023}$	$301.5^{+6.7}_{-6.4}$	$6.32^{+0.06}_{-0.06}$	$3.51^{+0.47}_{-0.48}$	$112^{+25}_{-23}$	105.9/118	0.90
Jul06	$18.8^{+0.4}_{-0.4}$	$1.844^{+0.013}_{-0.013}$	$544.5^{+6.8}_{-6.9}$	$6.34^{+0.05}_{-0.05}$	$4.21^{+0.51}_{-0.50}$	$77^{+15}_{-14}$	107.3/118	0.91
Dec06 <sup>h</sup>	$18.7^{+0.3}_{-0.3}$	$1.832^{+0.012}_{-0.011}$	$402.7^{+4.5}_{-4.6}$	$6.37^{+0.03}_{-0.04}$	$4.46^{+0.33}_{-0.33}$	$109^{+13}_{-13}$	134.4/117	1.15
Aug07 <sup>h</sup>	$17.3^{+0.2}_{-0.3}$	$1.818^{+0.009}_{-0.009}$	$736.4^{+6.5}_{-6.4}$	$6.30^{+0.05}_{-0.05}$	$4.27^{+0.51}_{-0.50}$	$56^{+11}_{-11}$	129.2/117	1.10
Jan09	$17.0^{+0.2}_{-0.1}$	$1.826^{+0.005}_{-0.006}$	$985.8^{+5.1}_{-5.0}$	$6.29^{+0.03}_{-0.04}$	$5.01^{+0.40}_{-0.40}$	$49^{+6}_{-7}$	137.1/118	1.16 <sup>j</sup>
Feb09	$17.9^{+0.2}_{-0.2}$	$1.830^{+0.008}_{-0.007}$	$992.3^{+7.5}_{-7.5}$	$6.28^{+0.04}_{-0.05}$	$4.81^{+0.59}_{-0.59}$	$46^{+10}_{-8}$	134.2/118	1.14

<sup>a</sup>Inferred Line of Sight Equivalent Hydrogen Column Density ( $10^{22} \text{ cm}^{-2}$ )

<sup>b</sup>Power Law Photon Index

<sup>c</sup>Power Law Normalization (2-10 keV;  $10^{-12} \text{ ergs cm}^{-2} \text{ s}^{-1}$ )

<sup>d</sup>The Line Energy (keV)

<sup>e</sup>The Line Flux ( $10^{-4} \text{ Photons cm}^{-2} \text{ s}^{-1}$ )

<sup>f</sup>The Line Equivalent Width (eV); **Uncertainties are 99% confidence**

<sup>g</sup>Total Chi-squared/Degrees of Freedom

<sup>h</sup>Additional systematics line at 29.5 keV included in fit

<sup>i</sup>Additional systematics lines at 4.5 and 29.5 keV included in fit

<sup>j</sup>A single PCU2 bin contributed  $\chi^2=21$ ; without it,  $\chi^2=113.1/117=0.97$

Table 3. Fluxes and Limits on Reflection and Cutoffs for the 13 Observations of Cen A with *RXTE*

Date	Flux(1 keV) <sup>a</sup>	Flux(2-10 keV) <sup>b</sup>	Flux(20-100keV) <sup>b</sup>	$\Gamma$	$E_{cut}$ <sup>c</sup>	$\chi^2_{\nu}$	R <sup>d</sup>	$\chi^2_{\nu}$
Aug96	0.099 <sup>+0.007</sup> <sub>-0.009</sub>	1.857 <sup>+0.007</sup> <sub>-0.007</sub>	4.99 <sup>+0.19</sup> <sub>-0.18</sub>	1.759 <sup>+0.058</sup> <sub>-0.092</sub>	154 <sup>+539</sup> <sub>-86</sub>	1.12	≤8.7	1.16
Aug98	0.092 <sup>+0.002</sup> <sub>-0.003</sub>	1.773 <sup>+0.003</sup> <sub>-0.002</sub>	4.54 <sup>+0.06</sup> <sub>-0.06</sub>	1.804 <sup>+0.015</sup> <sub>-0.027</sub>	≥204	1.17	≤2.1	1.20
Jan00	0.155 <sup>+0.005</sup> <sub>-0.005</sub>	2.843 <sup>+0.005</sup> <sub>-0.006</sub>	7.52 <sup>+0.08</sup> <sub>-0.11</sub>	1.832 <sup>+0.014</sup> <sub>-0.018</sub>	≥689	1.01	≤7.9	1.01
Mar03	0.203 <sup>+0.004</sup> <sub>-0.003</sub>	3.351 <sup>+0.002</sup> <sub>-0.003</sub>	11.23 <sup>+0.07</sup> <sub>-0.07</sub>	1.796 <sup>+0.007</sup> <sub>-0.006</sub>	≥1399	1.03	≤4.9	1.01
Jan04	0.120 <sup>+0.002</sup> <sub>-0.003</sub>	1.909 <sup>+0.002</sup> <sub>-0.002</sub>	7.01 <sup>+0.08</sup> <sub>-0.05</sub>	1.765 <sup>+0.015</sup> <sub>-0.015</sub>	607 <sup>+747</sup> <sub>-225</sub>	1.00	≤1.0	1.27
Feb04	0.136 <sup>+0.005</sup> <sub>-0.004</sub>	2.207 <sup>+0.004</sup> <sub>-0.003</sub>	7.63 <sup>+0.07</sup> <sub>-0.09</sub>	1.793 <sup>+0.013</sup> <sub>-0.021</sub>	≥693	1.07	≤4.4	1.08
Aug05	0.126 <sup>+0.005</sup> <sub>-0.008</sub>	2.048 <sup>+0.007</sup> <sub>-0.007</sub>	5.78 <sup>+0.14</sup> <sub>-0.12</sub>	1.845 <sup>+0.031</sup> <sub>-0.049</sub>	≥318	0.82	≤18.2	0.82
Dec05	0.090 <sup>+0.004</sup> <sub>-0.005</sub>	1.579 <sup>+0.005</sup> <sub>-0.004</sub>	4.55 <sup>+0.11</sup> <sub>-0.11</sub>	1.816 <sup>+0.027</sup> <sub>-0.049</sub>	≥209	0.92	≤12.6	0.91
Jul06	0.167 <sup>+0.005</sup> <sub>-0.005</sub>	2.767 <sup>+0.003</sup> <sub>-0.004</sub>	7.79 <sup>+0.11</sup> <sub>-0.12</sub>	1.842 <sup>+0.012</sup> <sub>-0.016</sub>	≥814	0.92	≤8.2	0.92
Dec06	0.121 <sup>+0.004</sup> <sub>-0.003</sub>	2.071 <sup>+0.002</sup> <sub>-0.003</sub>	5.89 <sup>+0.10</sup> <sub>-0.09</sub>	1.809 <sup>+0.026</sup> <sub>-0.025</sub>	≥242	1.12	≤7.1	1.14
Aug07	0.216 <sup>+0.005</sup> <sub>-0.004</sub>	3.887 <sup>+0.004</sup> <sub>-0.003</sub>	11.11 <sup>+0.10</sup> <sub>-0.11</sub>	1.815 <sup>+0.009</sup> <sub>-0.013</sub>	≥878	1.11	≤5.7	1.10
Jan09	0.293 <sup>+0.004</sup> <sub>-0.003</sub>	5.236 <sup>+0.003</sup> <sub>-0.003</sub>	14.68 <sup>+0.10</sup> <sub>-0.09</sub>	1.823 <sup>+0.006</sup> <sub>-0.004</sub>	≥2796	1.21	≤4.6	1.16
Feb09	0.297 <sup>+0.006</sup> <sub>-0.005</sub>	5.139 <sup>+0.003</sup> <sub>-0.005</sub>	14.60 <sup>+0.12</sup> <sub>-0.11</sub>	1.828 <sup>+0.008</sup> <sub>-0.008</sub>	≥2499	1.17	≤3.9	1.15

<sup>a</sup>Power law normalization at 1 keV (photons cm<sup>-2</sup> s<sup>-1</sup> keV<sup>-1</sup>)

<sup>b</sup>Flux in units of 10<sup>-10</sup> ergs cm<sup>-2</sup> s<sup>-1</sup>

<sup>c</sup>Exponential Cut-off Energy (keV)

<sup>d</sup>Compton Reflection Strength (%)

Table 4. Start/Stop Times for the 51 PCU2 Light Curves of Cen A

Obs. Interval	Date	Time (UT)	Obs. Interval	Date	Time (UT)	Obs. Interval	Date	Time (UT)
August 1996	8/14	2:43-8:03	December 2005	12/16	0:08-0:21	August 2007	8/18	3:54-23:27
August 1998	8/9	1:24-23:56		12/17	17:06-22:09		8/19	0:19-15:03
	8/10	0:01-7:18		12/20	20:00-24:07		8/20	14:10-24:07
	8/14	3:08-23:58		12/23	0:08-22:52		8/22	3:44-7:51
	8/15	0:02-8:57	July 2006	7/13	10:08-24:27	January 2009	1/15	20:34-23:58
January 2000	1/23	7:38-23:24		7/14	11:03-22:28		1/16	0:02-19:20
March 2003	3/7	12:02-20:58		7/16	17:56-23:04		1/18	22:24-23:58
	3/8	1:04-17:28		7/17	0:13-17:43		1/19	0:02-23:58
	3/9	3:33-18:48		7/18	10:54-12:38		1/20	0:02-1:57
	3/10	8:00-16:02		7/19	13:32-16:54		1/25	22:19-23:58
	3/11	7:43-8:17		7/20	11:39-22:55		1/26	0:02-0:44
January 2004	1/2	8:56-23:58	December 2006	12/8	7:03-13:45		1/27	18:16-23:49
	1/3	0:03-23:47		12.9	1:39-2:21		1/28	0:32-1:23
	1/4	0:46-16:29		12/10	1:47-9:26	February 2009	2/2	13:58-22:39
February 2004	2/13	9:41-16:46		12/11	4:26-16:19		2/20	4:59-19:19
	2/14	9:19-19:39		12/12	9:39-13:12		2/21	4:06-7:53
August 2005	8/20	3:59-7:58		12/13	9:09-19:14			
				12/15	5:10-18:19			