

Remote Raman Spectroscopy of Minerals at Elevated Temperature Relevant to Venus Exploration

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ABSTRACT

We have used a remote time-resolved telescopic Raman system equipped with 532 nm pulsed laser excitation and a gated intensified CCD (ICCD) detector for measuring Raman spectra of a number of minerals at high temperature to 970 K. Remote Raman measurements were made with samples at 9-meter in side a high-temperature furnace by gating the ICCD detector with 2 micro-sec gate to minimize interference from blackbody emission from mineral surfaces at high temperature as well as interference from ambient light. A comparison of Raman spectra of gypsum ($\text{CaSO}_4 \cdot 2\text{H}_2\text{O}$), dolomite ($\text{CaMg}(\text{CO}_3)_2$), and olivine ($\text{Mg}_2\text{Fe}_{2-x}\text{SiO}_4$), as a function of temperature shows that the Raman lines remains sharp and well defined even in the high-temperature spectra. In the case of gypsum, Raman spectral fingerprints of $\text{CaSO}_4 \cdot \text{H}_2\text{O}$ at 518 K were observed due to dehydration of gypsum. In the case of dolomite, partial mineral dissociation was observed at 973 K at ambient pressure indicating that some of the dolomite might survive on Venus surface that is at ~ 750 K and 92 atmospheric pressure. Time-resolved Raman spectra of low clino-enstatite (MgSiO_3) measured at 75 mm from the sample in side the high-temperature furnace also show that the Raman lines remains sharp and well defined in the high temperature spectra. These high-temperature remote Raman spectra of minerals show that time-resolved Raman spectroscopy can be used as a potential tool for exploring Venus surface mineralogy at shorter (75 mm) and long (9 m) distances from the samples both during daytime and nighttime. The remote Raman system could also be used for measuring profiles of molecular species in the dense Venus atmosphere during descent as well as on the surface.

Keywords: Venus, telescopic Raman instrument, time-resolved Raman spectroscopy, high-temperature Raman spectra, mineralogy, Raman spectra of molecular gases

1. INTRODUCTION

Raman spectroscopy (RS) has been proposed as a candidate analysis system for missions to Mars, and more recently as a potential mineralogical analysis system to Venus. This technique has several distinct advantages over other spectroscopic techniques that have been used on past missions. For planetary applications, the sharpness of Raman spectral features of minerals allows for much less ambiguous detection, especially in the presence of mixtures. Visible, near-infrared, thermal, and reflectance and emission spectroscopy of minerals and compounds all suffer from broad overlapping spectral features [1] which complicate interpretation of their spectra. Small portable remote Raman systems, which are suitable for planetary rovers and landers, have been shown to be effective in identifying hydrous and anhydrous minerals, glasses of mineral compositions, and ices. The remote Raman systems developed at the University of Hawaii have been tested from 9 m to ≥ 100 m distance [e.g., Refs. 2–7]. Two of the most important advantages of time-resolved (TR) RS have over other techniques for a mission to Venus is its rapid mineralogical analysis of both hydrous and anhydrous minerals and stand-off analysis at distances of many meters. Rapid mineralogical analysis and stand-off analysis are very important to missions to Venus due to the harsh environment at the planet surface. From the Venera missions it is known that on the Venusian surface the pressures are approximately 9.1 MPa (90 atm) and the temperature is near 735 K [8]. For these reasons, the Soviet Venera surface probes had operational lifetimes of less than 2 hours. The major gas in Venus atmosphere is CO_2 , and under Venusian atmospheric condition CO_2 exists in a supercritical phase near the surface. In 2003, the National Research Council recommended in its Solar System Exploration Decadal Survey [9] that “NASA commit to significant new investments in advanced technology so that

future high-priority flight missions can succeed.” These proposed high-priority flight missions were further discussed in NASA’s 2006 Solar System Exploration Roadmap [10]. Currently Venus is under consideration for NASA’s New Frontier Program. The Venus Exploration and Analysis Group (VEXAG) has recommended that NASA should include Venus In-situ Explorer (VISE) in Announcement of Opportunities (AO) of New Frontiers Class mission for 2008 [11]. The Explorer will be investigating the composition and surface properties of Venus. It will also acquire and characterize a core sample from the surface of Venus, and it will measure the elements and mineralogy of surface materials. VEXAG has also recommended that NASA should initiate a study of a Venus Flagship Mission at earliest opportunity [11]. In light of these future NASA missions, it is beneficial to evaluate different analysis methods such as remote TR Raman spectroscopy, which offer to greatly increase the scientific return from such a mission.

Currently we have begun to evaluate TR remote Raman spectroscopic detection in an environment with temperatures similar to those found on Venus. Although the pressure of Venus (91 atm) has not been taken into account in these experiments, it is anticipated that it will not affect Raman analytical capability. Blackbody radiation at high temperature could interfere with Raman spectroscopic measurements. Time-resolved Raman spectroscopy with pulsed 532 nm laser has been shown to be an effective technique for high-temperature Raman spectroscopy [12]. Previous work has shown that the high temperatures can have a strong effect on the Raman spectra of minerals [12-13]. These temperature effects include broadening of Raman lines, decrease in the intensity of Stokes-Raman lines, phase transitions in minerals, dehydration and de-carbonation of minerals. High-pressure and high-temperature studies have been conducted on minerals at very high-pressures —10’s to 50’s GPa in diamond anvil cells (e.g., [12-13]). The results of these investigations show that the Raman spectra of minerals at Venus surface pressure of 90 atm will only show small changes in the positions of Raman lines. However, dehydration and de-carbonation of minerals could be affected significantly by high atmospheric pressure of CO₂ on Venus. Here we show results from a remote Raman study of minerals to 973 K, and of CO₂ gas at 1 atm at 9 m, which demonstrates the feasibility of using this technique for mineralogical analyses and atmospheric studies at Venusian surface.

2. INSTRUMENTATION AND EXPERIMENTAL METHOD

Time-resolved Remote Raman system

In a conventional Raman system utilizing a CW laser, a sample is excited with a continuous beam of photons but only a very small fraction of these incident photons generate Raman signals, because of low Raman cross-section. In such systems, Raman spectra are collected by integrating the signal over several seconds under dark conditions, which helps in maintaining low background noise. When exposed to daylight, these systems become unusable, because the detector becomes oversaturated from the high background light level. Because of high surface temperature (~735 K) on Venus surface, the thermal blackbody radiation from the sample will also interfere with the Raman spectra of minerals.

One solution to this problem is to generate large numbers of Raman photons in a relatively short time and then detect them more efficiently. A remote Raman system for daytime operation has two critical components: (1) a pulsed laser and (2) a gated detector. A time-resolved Raman system significantly improves the Raman signal-to-background noise ratio and successfully measures Raman spectra in daylight as well as from the samples that are at high temperature. Using a laser with pulse widths of a few nanoseconds, the target is excited with large numbers of photons within the nanosecond time frame. The Raman phenomenon decays quickly, producing a signal, which lasts only a few nanoseconds after the laser pulse exposure. If the detector is only activated within this small time window, the number of Raman photons is large in comparison to the background photons, resulting in high signal-to-background ratio. In general, the number of Raman photons generated is directly proportional to the pulse energy of the laser. Likewise, the amount of background light detected by the system is directly proportional to the gate width of the detector. Therefore, precise time-gating of the detection system is critical in capturing a maximum number of Raman photons and minimum background light energy.

The directly coupled and portable remote pulsed Raman system used in the current investigations is described in detail elsewhere [4-7], and a schematic diagram of the system is shown in Fig. 1. The remote time-resolved Raman system consists of an 8-inch (203.2 mm) diameter telescope (Meade LX-200R Advanced Ritchey-Chretien, f/10), a frequency-doubled mini Nd:YAG pulsed laser source (model Ultra CFR, Big Sky Laser, 532 nm, 35 mJ/pulse, 20 Hz, pulse width 8 ns, central laser spot divergence ~0.5 mrad), and a Kaiser Optical Systems f/2.2 HoloSpec spectrometer equipped with a gated, thermo-electrically cooled ICCD detector. The telescope is directly coupled to the spectrometer through a camera lens (50 mm, f/1.8). A 532-nm holographic SuperNotch filter (SNF) is placed in front of the camera lens to minimize the reflected and Rayleigh scattered photons from the target. The signal is then directed through a fixed mechanical slit of

100 μm width into the spectrograph. The entire system is mounted atop a pan-and-tilt scanner, which is electronically controlled, on a portable trolley. Since the entire system moves together, optical alignment between the laser, telescope, and spectrograph is always maintained when the system is configured in the co-axial mode. A 10x beam expander is mounted in front of the laser to allow focusing of the laser spot at targeted distances. After focusing, the laser spot size is typically 1 mm in diameter at 50 m distance. For Raman measurements, the laser beam diameter at the sample is kept to 7 mm. The laser is operated at 35mJ/pulse with a frequency of 20 Hz. The HoloSpec spectrograph is equipped with a volume phase holographic grating capable of measuring Stokes-shifted Raman spectra in the range of 50 to 2500 cm^{-1} . For some analyses, this grating was replaced by a HoloPlex grating (Kaiser Optical Systems) which extends the spectral range to 4500 cm^{-1} while maintaining a resolution of 7 cm^{-1} .

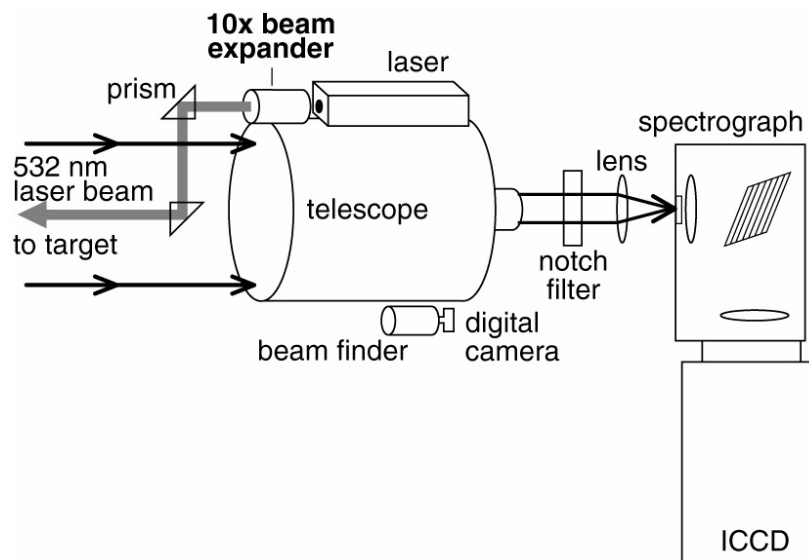


Fig. 1. Schematic diagrams of directly coupled remote time-resolved Raman spectrograph systems shown in coaxial geometry.

All spectra presented here were measured with the ICCD detector operating in the gated mode. For timing purposes, trigger signals generated by the flash lamp of the laser were used as a reference. The gate delay of the ICCD detector was adjusted so that it opened at the same time as the arrival of the first Raman signal. In general, since the pulse width of our Nd:YAG laser is 8 ns, the optimum detector gate width is expected to be only a few nanoseconds, during which Raman photons are arriving. However, due to slight fluctuations of the electronics in the laser, laser pulses are not consistently generated with the exact same time interval. In this study, fluctuations in the pulse generation are on the order of $\pm 1 \mu\text{s}$. Hence, a gate width of 2 μs was used to capture Raman signals excited by the laser pulses. Samples were also measured with an integration time of 1 s, which is equivalent to exposure of 20 laser pulses for our system operating at 20 Hz with pulse energy of 35mJ/pulse.

No cosmic ray or baseline corrections were applied to the spectra presented in this study. This is an additional advantage of using the gated mode: the ICCD acquires the spectrum only during the interval for which the gate is open, and a gate width of 2 μs significantly reduces the probability of capturing a cosmic ray. Thus, most spectra are free from cosmic peaks.

With short-distance ($\sim 10\text{m}$) remote Raman systems, we commonly use the atomic emission lines of a neon lamp to calibrate our spectrometer. In addition, liquid cyclohexane and atmospheric nitrogen (band at 2331 cm^{-1}) were used as Raman standards to calibrate the system (cyclohexane peak positions were obtained from the NIST website [14]). The Raman peak positions of various minerals reported here are within $\pm 2 \text{ cm}^{-1}$ of their standard literature values. Two commercial computer programs were used for data collection and spectral processing: WinSpec®, a 32-bit Windows software package (Roper Scientific) for spectral data acquisition, and GRAMS/32® software package (Galactic Industries) for data processing.

High Temperature Furnace

Mineral samples were heated to high temperature in a windowless furnace at 9 m from the telescope show in Fig.2. The furnace is heated with a platinum-10% Rhodium wire and a temperature up to 1400 K can be achieved. The mineral is placed inside the furnace close to the center on a ceramic tube, and the laser beam is focused with the 10x beam expander mounted on the top of the telescope. The backscattered light from the sample is collected by the telescope and focused on the slit of the gated Raman spectrometer.

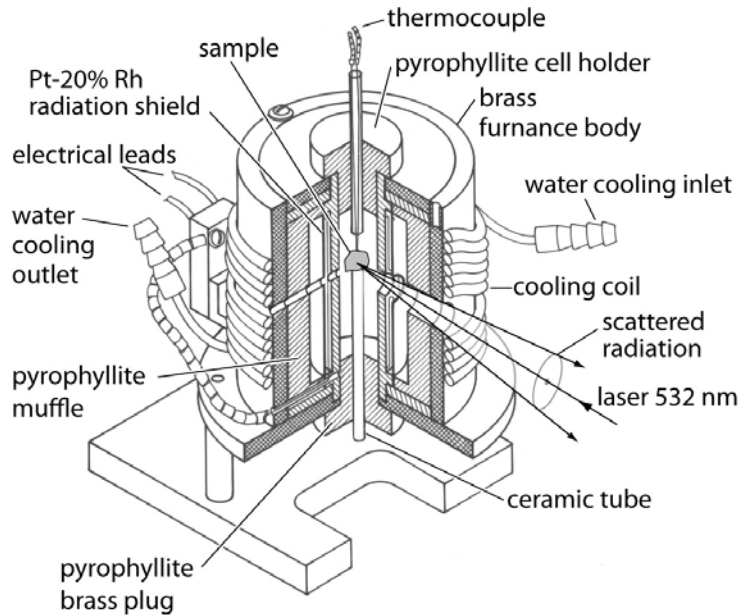


Fig.2. A windowless Pt-10%Rh wire-wound furnace for measuring Raman spectra of minerals at high temperature.

Mineral Samples

The rock-forming minerals (dolomite, gypsum and dunite rock) were purchased from Ward's Natural Science Establishment, Inc., Rochester, New York. The polycrystalline dolomite sample was from Bamble, Norway. Gypsum of the selenite variety was from Washington County, Utah, USA. Dunite rock contained 90% fine grain olivine crystal and was from Balsam, North Carolina, USA. The low clino-enstatite crystals were provided by Professor Subrata Ghose of University of Washington. A small fragment of these samples were mounted inside the furnace and was excited with

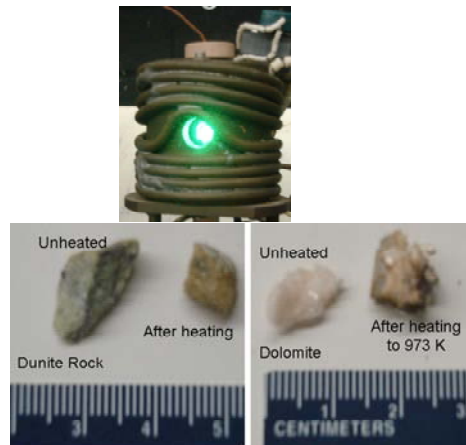


Fig. 3. Photographs of the high-temperature furnace and of dunite and dolomite sample before and after heating. Furnace is not to scale.

532-nm pulsed laser without further cleaning or polishing. Figure 3 shows the high-temperature furnace and the samples before and after heating for Raman measurements. The dunite sample changes color because of oxidation of olivine in the air at high temperature.

3. RESULTS AND DISCUSSION

Time-resolved Raman spectra of gypsum in the ranges 100-2400 cm^{-1} , and 3000-3800 cm^{-1} at 9 m and at 303, 373 and 518 K are shown in Figs. 4 and 5, respectively. The line widths of the strongest Raman line at 1008 cm^{-1} of $\nu_1(\text{SO}_4)$,

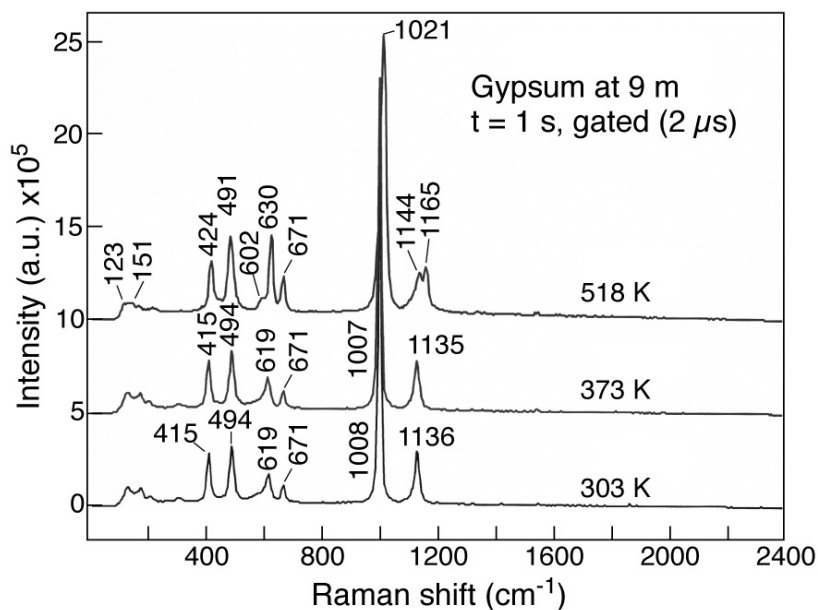


Fig. 4. Time-resolved high-temperature Raman spectra of gypsum at 9 m in the 100-2400 cm^{-1} region. Laser 532 nm, 20 Hz, 35 mJ/pulse, averaged over 20 pulses (1 s)

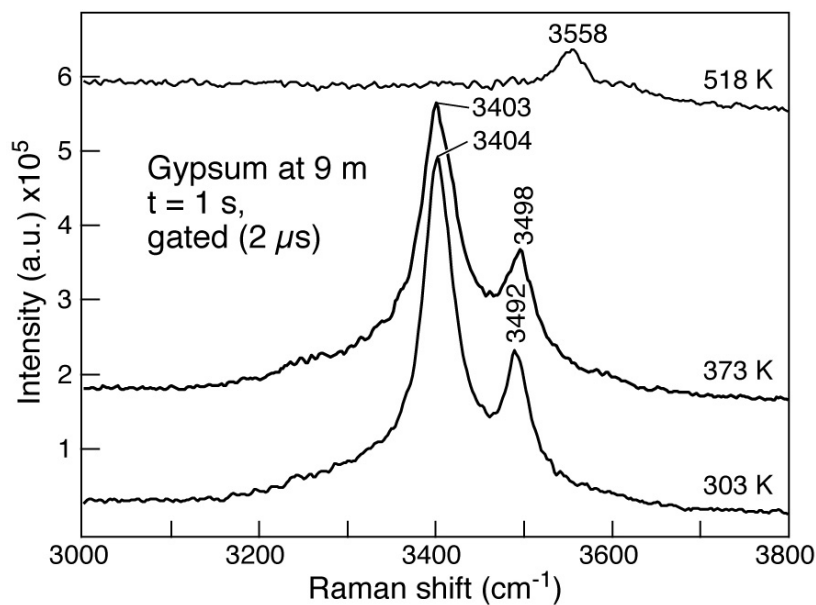


Fig. 5. Time-resolved high-temperature Raman spectra of gypsum at 9 m in the 3000-3800 cm^{-1} region. Laser 532 nm, 20 Hz, 35mJ/pulse, averaged over 20 pulses (1 s).

symmetric stretching mode of sulfate ions in gypsum at various temperatures are listed in Table 1. It is evident from Figs. 4 and 5 that the Raman spectra of gypsum can be measured at high temperatures both in the 100-1300 cm^{-1} and in the O-H stretching region with the time-resolved remote Raman system. On increasing temperature from 303 K to 373 K, the Raman lines of gypsum in 100-1300 cm^{-1} region show shift to lower wave number and their line widths increase. The peak position of the H_2O (I) stretching mode at 3404 cm^{-1} observed at 303 K decreases to 3403 cm^{-1} at 373 K where as the H_2O (II) stretching mode of water molecules shows an increase in the peak position to 3498 cm^{-1} at 373 K (Table 1). These results are consistent with previous high-temperature micro-Raman laboratory studies of gypsum [15-16]. On heating the gypsum sample to 518 K, the $\nu_1(\text{SO}_4)$ peak appears at 1021 cm^{-1} with line width of 24 cm^{-1} , the $\nu_3(\text{SO}_4)$ anti-symmetric stretching mode of sulfate appears as a doublet (1144 and 1165 cm^{-1}), and in the O-H stretching region a band appears at 3558 cm^{-1} . These changes suggest that there is a phase transition in gypsum. Previous studies have shown that on heating gypsum above 373 K causes gypsum to lose water molecules during the heating process forming $\gamma\text{-CaSO}_4$ (soluble form of anhydrite) and hemihydrate ($\text{CaSO}_4 \cdot 0.5\text{H}_2\text{O}$ also known as bassanite) [15, 17-18]. The Raman $\nu_1(\text{SO}_4)$ lines of hemihydrate and $\gamma\text{-CaSO}_4$ appear at 1014 and 1026 cm^{-1} , respectively. The large (24 cm^{-1}) line width of the Raman line at 1021 cm^{-1} , the presence of 1014 and 1026 cm^{-1} doublet and the band at 3558 cm^{-1} in the 518 K spectrum indicate that both hemihydrate and $\gamma\text{-CaSO}_4$ phases are present in the heated sample. Heating of gypsum to 633 K result in the formation of $\alpha\text{-CaSO}_4$ (anhydrite) [15].

As the Venus surface temperature is 735 K, hydrous phases of CaSO_4 will not be stable and will convert to anhydrite phase. The high-temperature remote Raman measurements on gypsum do exhibit the potential of the technique for searching hydrous phases and their thermal decomposition products on the Venus' surface from a lander. On the basis of high D/H ratio (150 times the terrestrial value) in the Venus lower atmosphere, it has been suggested that Venus may have once had a greater abundance of water [19, 20]. If so, the detection of hydrous minerals on Venus would prove that Venus had more water in the past. Laboratory investigations on hydrous amphibole minerals, such as tremolite, under Venus-like conditions have shown that these minerals if formed on Venus surface during a wetter past could still exist at current conditions on Venus' surface [21].

Table 1. Raman peak positions and line widths (in cm^{-1}) of major lines of gypsum, dolomite and olivine at various temperatures

Minerals	T=303 K	T=373 K	T=518K
Gypsum	1008 (ω *=13.3)	1007 (ω =13.6)	1021 (ω =25)
	3401 (ω =40.5)	3400 (ω =53)	
	3489 (ω =41)	3496 (ω =56)	
			3555 (ω =36)
Dolomite	T=303 K	T=723 K	T= 973 K
	1094 (ω =14.7)	1087.5 (ω =22.2)	1081 (ω =23.8)
Olivine	T=303 K	T= 718 K	T= 878 K
	824 (ω =18.41)	816 (ω =24.5)	813 (ω =29.1)
	855 (ω =19.6)	850 (ω =25.4)	847 (ω =26.6)

ω = Full line width at half-maximum height in cm^{-1}

High-temperature time-resolved Raman spectra of dolomite at 9 m distance are shown in Fig. 6. The line widths of the 1094 cm^{-1} Raman line at various temperatures are listed in Table 1. Raman peaks attributed to the lattice and internal modes of dolomite are observed at 174, 297, 723, 1094, 1439, and 1752 cm^{-1} [22, 23]. The strongest 1094 cm^{-1} Raman band is the fingerprint of carbonate ions corresponding to the $\nu_1(\text{CO}_3)$ symmetric stretching mode of carbonate oxygen atoms. The low-frequency Raman bands at 174 and 297 cm^{-1} are characteristic of dolomite structure. The 723 cm^{-1}

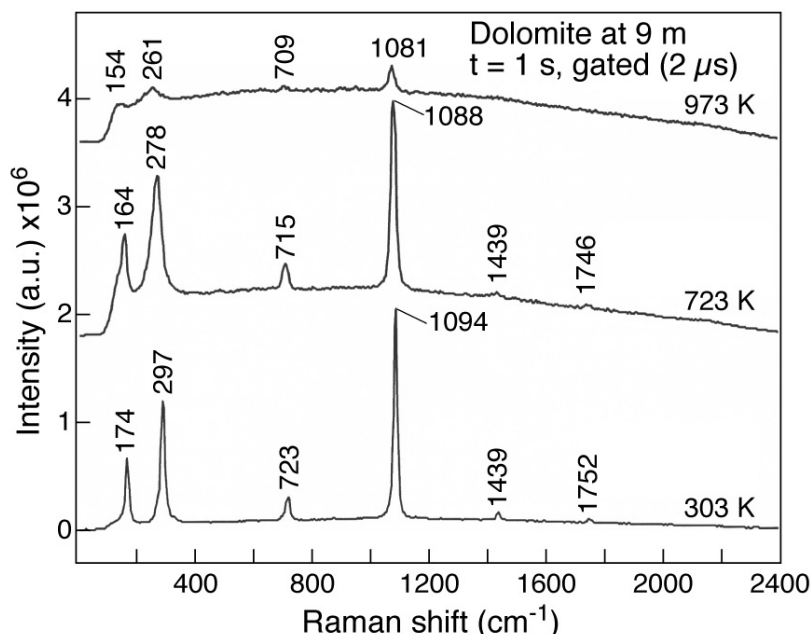


Fig. 6. Time-resolved high-temperature Raman spectra of dolomite at 9 m in the 100-2400 cm^{-1} region. Laser 532 nm. 35 mJ/pulse. averaged over 20 pulses (1 s).

band corresponds to the $\nu_4(\text{CO}_3)$ in-plane bending mode, and the very weak band at 1439 cm^{-1} corresponds to the $\nu_3(\text{CO}_3)$ anti-symmetric stretching mode of carbonate ions. Another weak band at 1752 cm^{-1} is an overtone mode of the infrared-active $2\nu_2(A_g)$ mode [22]. In addition, it is possible to distinguish various minerals within the same functional group from each other by the positions of the low-frequency lattice modes and the mid-frequency internal modes [2-3, 22-25]. The lattice modes of dolomite (CaMgCO_3) ($174, 297 \text{ cm}^{-1}$), calcite (156 and 282 cm^{-1}) and magnesite (MgCO_3) (212 and 329 cm^{-1}) are widely separated and easily identifiable. Because of their unique Raman patterns and clear, distinguishable peaks, identification of minerals and inorganic compounds by their Raman spectra becomes promising. In fact, Raman spectroscopy can be used as an unambiguous analytical tool for identification of a mineral of interest by comparing its spectra with the library of Raman patterns.

With increasing temperature to 723 K the Raman lines of dolomite in the time-resolved remote Raman spectrum (Fig. 6) shift to lower frequencies and their widths increase. At 973 K the Raman spectrum of dolomite shows sudden decrease in the line intensities and there are marked changes in the low-frequency lattice mode region. These changes in the 973 K Raman spectrum of dolomite indicate phase transition in the material. Differential thermal analysis and high temperature X-ray diffraction studies have shown that on heating dolomite dissociates into CaCO_3 and MgO between 833 and 1038 K [26]. MgO is a cubic crystal and therefore in the bulk has no first-order Raman spectrum as it possesses inversion symmetry. It is, therefore, not possible to detect MgO from the Raman spectrum [27]. The changes in the structure of dolomite on heating to 973 K are, however, observed in the spectrum of the carbonate phase. The Raman spectrum of dolomite at 973 K indicates that the sample is an intermediate phase in the thermal decomposition process and partly resembles that of calcite spectrum. These high-temperature Raman measurements suggest that the carbonates could be a stable phase on Venus surface, and could be detected with the time-resolved remote Raman system.

Time-resolved Raman spectra of olivine in a dunite rock, which contains 90% fine grain olivine crystals, at 303, 718 and 878 K at 9 m are shown in Fig. 7. These spectra were recorded by integrating the spectra over 200 laser shots. The peak frequency and line widths of the most prominent Raman lines in the $800\text{-}850 \text{ cm}^{-1}$ region at various temperatures are listed in Table 1. The intensities of Raman lines of olivine are comparable to those of atmospheric O_2 (1556 cm^{-1}) and N_2 (2331 cm^{-1}) indicating that olivine is a weaker Raman scatterer than gypsum and dolomite [25]. It was, however, possible to measure Raman spectra of olivine to 878 K at 9 m with good signal-to-noise ratio within 10 s. These high-temperature Raman spectra are in agreement with the results of high-temperature laboratory Raman spectroscopic studies of olivine crystals [27, 28].

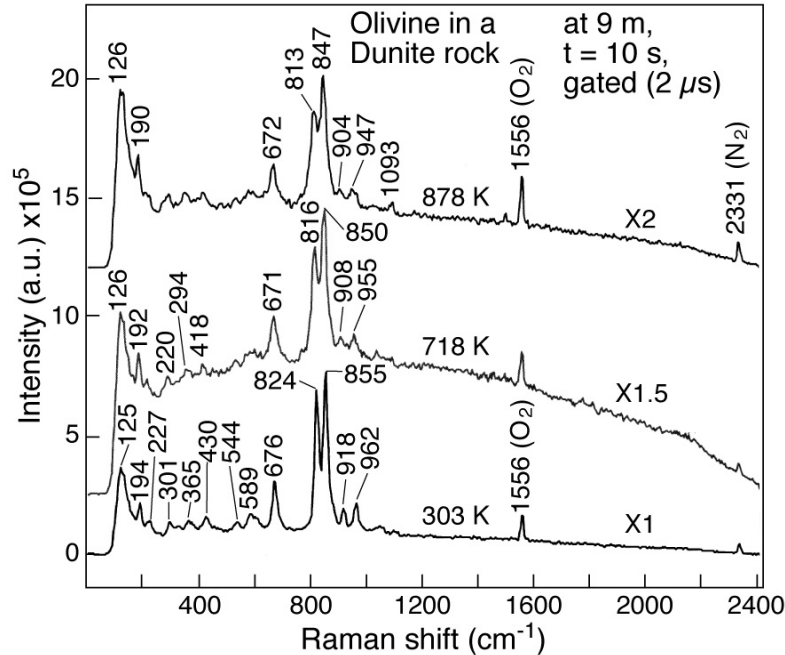


Fig. 7. Time-resolved high-temperature Raman spectra of olivine in a dunite rock at 9 m in the 100-2400 cm^{-1} region.

The remote Raman spectrum of olivine at 303 K shows lines at 125, 194, 227, 365, 430, 544, 589, 676, 824, 855, 918, and 962 cm^{-1} . The 824 and 855 cm^{-1} doublet in the Raman spectrum of olivine at 303 K is the Raman fingerprint of olivine.

Forsterite (Fo, Mg_2SiO_4) and fayalite (Fay, Fe_2SiO_4) are the end member minerals of olivine group. In forsterite at room temperature the signature double peaks are observed at 825 and 857 cm^{-1} . The corresponding Raman lines of fayalite at room temperature appear at 812 and 835 cm^{-1} . These Raman lines are due to the ν_1 - and ν_3 -derived vibrational modes of orthosilicate (SiO_4) units. The ν_1 and ν_3 , respectively, correspond to Si-O_{nb} symmetric and anti symmetric stretching modes of the SiO_4^{4-} tetrahedra, where O_{nb} refers to non-bridging oxygen. The frequencies of these ν_1 - and ν_3 -derived modes show small variation with substitution of Fe or Ca for Mg [28-33]. Presence of Raman lines at 824 and 855 cm^{-1} indicate that the dunite rock contains Fo-rich olivine crystals.

On increasing the temperature of olivine crystals in dunite rock to 718 K the Raman lines shift to lower frequency and their band widths increase (Table 1). This trend continues to 878 K, a temperature that is higher than the Venus surface temperature of ~ 735 K. In the olivine spectrum at 878 K, the line widths of 813 and 847 cm^{-1} Raman line are, respectively, 29.1 and 26.6 cm^{-1} . Evidently time-resolved remote Raman system could easily detect olivine on the Venus surface.

Figure 8 depicts time-resolved Raman spectra of low clinoenstatite measured at a distance of 75 mm from the collecting lens inside the furnace at 293, 373, 473, 673 and 873 K [11]. In the spectrum at 293 K, the $\nu_s(\text{Si-O-Si})$ modes of silicate chains appear at 668.6 and 691.5 cm^{-1} , and the $\nu_s(\text{Si-O}_{\text{nbr}})$ modes of silicon-non-bridging oxygen stretch of the pyroxene chain appear at 1007.3 and 1027.3 cm^{-1} .

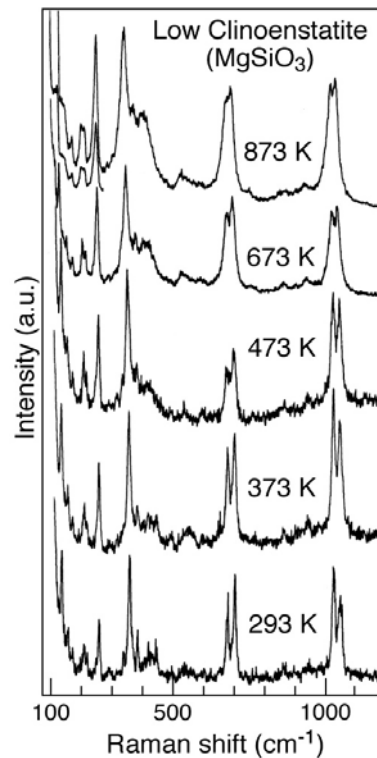


Fig. 8. Time-resolved Raman spectra of clinoenstatite at high-temperatures. (Sample at 75 mm from collection optics in side the furnace)

Table 2. Peak positions and line width (in cm^{-1}) of strong Raman lines of low clinoenstatite at selected temperatures

293 K	473 K	873 K	Assignments
668.6 ($\omega^*=10.7$)	665.7 ($\omega=17.9$)	663 ($\omega=20.7$)	$\nu_s(\text{Si-O-Si})$
691.6 ($\omega=10.7$)	688.2 ($\omega=17.9$)	683.6 ($\omega=17.8$)	$\nu_s(\text{Si-O-Si})$
1015.2 ($\omega=12$)	1012.0 ($\omega=16$)	1007.3 ($\omega=20.6$)	$\nu_s(\text{Si-O}_{\text{nbr}})$
1037.4 ($\omega=15$)	1034.6 ($\omega=15.6$)	1027.3 ($\omega=18.3$)	$\nu_s(\text{Si-O}_{\text{nbr}})$

* ω = Full line width at half maximum height (in cm^{-1})

With increasing temperature all the Raman lines of clinoenstatite shift toward lower wave number and their line widths increase. The frequencies and the line width of the $\nu_s(\text{Si-O-Si})$ and $\nu_s(\text{Si-O}_{\text{nbr}})$ modes at selected temperatures are listed in Table 2. It is clear from Fig. 8 that even at 873 K the Raman spectrum of clinoenstatite has good signal to noise ratio, and the fingerprint lines of pyroxene corresponding to $\nu_s(\text{Si-O-Si})$ and $\nu_s(\text{Si-O}_{\text{nbr}})$ modes are clearly visible. Evidently, time-resolved Raman spectroscopic technique can be used on Venus surface for investigating mineralogy of silicate minerals both at short (75 mm) and at relatively long distance of 9 m. We have demonstrated that the time-resolved Raman spectroscopy can be used for investigating Venus surface mineralogy including silicates, carbonates, sulfates and hydrous and anhydrous minerals.

Figure 9 shows a time-resolved remote Raman spectrum of carbon dioxide (CO_2) at 1 atm in fish tank of 120 cm length at 9 m distance. In Raman spectroscopy, the ν_1 totally symmetric stretching vibration of CO_2 (1305.4 cm^{-1}) is active for the isolated molecule (e.g., the monomer) and is in Fermi resonance with the $2\nu_2$ overtone of the bending mode of the molecule ($\nu_2 \sim 667 \text{ cm}^{-1}$). As a consequence, two intense peaks are observed giving rise to the so-called Fermi doublet, situated at 1285 and 1388 cm^{-1} . The spectrum in Fig. 9 clearly shows a Fermi resonance doublet of CO_2 molecules at 1285 ($2\nu_2$) and 1388 (ν_1) cm^{-1} . Additional Raman lines (Fig. 9) at 575 and 1091 cm^{-1} are, respectively, $\nu_s(\text{Si-O-Si})$ symmetric stretching of bridging oxygen, and $\nu_s(\text{Si-O}_{\text{nbr}})$ stretching modes of non-bridging oxygen of the glass wall.

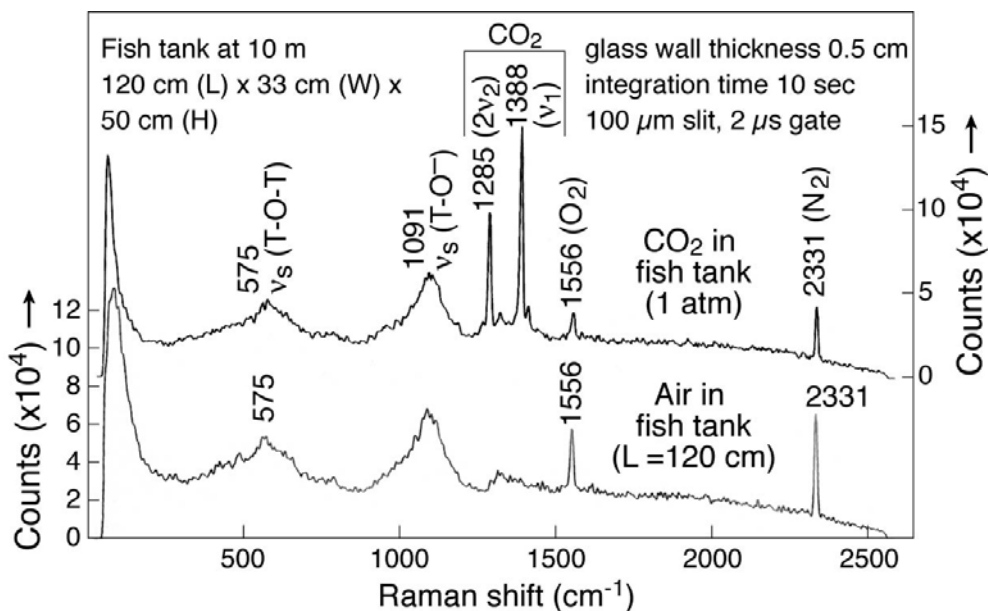
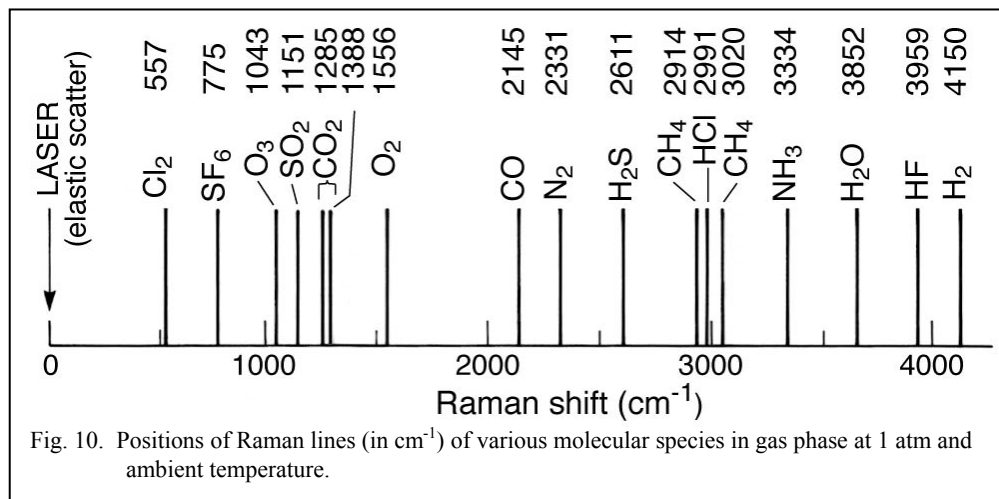


Fig. 9 Raman spectra of air and CO_2 in a fish tank at 9 m within a well illuminated laboratory. Laser 532 nm, 35 mJ/pulse, 20 Hz, spectra averaged over 200 laser shots.

The sharp lines at 1556 and 2331 cm^{-1} are, respectively, the symmetric stretching modes of atmospheric oxygen and nitrogen molecules. On the Venus surface, CO_2 pressure will be ~ 90 atm. The Raman spectrum will be dominated by

the 1285 and 1388 cm^{-1} lines. Laboratory Raman spectroscopic investigation of supercritical CO_2 in the pressure range of 0.1 to 35 MPa and at temperatures of 307, 309, 313, and 323 K have shown formation of transient CO_2 dimers as indicated by the presence of the ν_2 bending mode (forbidden transition at about 668 cm^{-1}) in the spectrum [34]. Figure 10 shows the positions of the strong Raman lines of various molecular species in the gas phase at 1 atm pressure and room temperature [35]. As the Raman lines of molecular gases are relatively sharp and will remain well defined at Venus atmospheric pressure and temperature, and will increase in intensity with increasing pressure, these lines can be used for profiling molecular species in Venus atmosphere. The elastic scattered light could provide profiles of particulates in Venus atmosphere by operating the Raman system in Mie-Rayleigh lidar mode [e.g., 24, 36].



4. CONCLUSIONS

We have demonstrated that it is feasible to measure good quality Raman spectra of minerals at high-temperature with a time-resolved remote Raman system during daytime and nighttimes. The results of the present investigation have also shown that Raman spectroscopy could provide information about the hydrous and anhydrous sulfate, carbonate and silicate minerals at high temperatures. The dehydration and decarbonation of minerals at high temperatures can be monitored from the reaction products of these minerals. In the case of dolomite it is found that mineral partially dissociate in to calcite and MgO . It seems that some of the carbonate minerals could survive on Venus surface under current conditions. The time-resolved Raman system is also capable of measuring molecular species in the atmosphere. The time-resolved Raman system could be very useful in future lander and rover missions to Venus for investigating surface mineralogy and profiling molecular species in Venus atmosphere during descending of the space crafts as well as on the surface.

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