

A Comprehensive Search for Gamma-Ray Lines in the First Year of Data from the INTEGRAL Spectrometer

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ABSTRACT

Gamma-ray lines are produced in nature by a variety of different physical processes. They can be valuable astrophysical diagnostics providing information that may be unobtainable by other means. We have carried out an extensive search for gamma-ray lines in the first year of public data from the Spectrometer (SPI) on the INTEGRAL mission. INTEGRAL has spent a large fraction of its observing time in the Galactic Plane with particular concentration in the Galactic Center (GC) region (~ 3 Msec in the first year). Hence the most sensitive search regions are in the Galactic Plane and Center. The phase space of the search spans the energy range 20-8000 keV, and line widths from 0-1000 keV (FWHM) and includes both diffuse and point-like emission. We have searched for variable emission on time scales down to ~ 1000 sec. Diffuse emission has been searched for on a range of different spatial scales from $\sim 20^\circ$ (the approximate field-of-view of the spectrometer) up to the entire Galactic Plane. Our search procedures were verified by the recovery of the known gamma-ray lines at 511 keV and 1809 keV at the appropriate intensities and significances. We find no evidence for any previously unknown gamma-ray lines. The upper limits range from a few $\times 10^{-5}$ $\text{cm}^{-2} \text{s}^{-1}$ to a few $\times 10^{-3}$ $\text{cm}^{-2} \text{s}^{-1}$ depending on line width, energy and exposure. Comparison is made between our results and various prior predictions of astrophysical lines.

Subject headings: gamma-rays: observations

1. Introduction

1.1. Astrophysical Gamma-Ray Line Production

Gamma-ray lines can be the bearers of important information on astrophysical sites and processes. A large body of work over the past 3 decades has identified a wide variety of

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different astrophysical scenarios in which gamma-ray line production can occur (for recent reviews see Knödlseeder & Vedrenne (2000); Prantzos (2004); Vink (2005)). The most important of these are summarized as follows:

Nucleosynthesis in Massive Stars: Radioactive elements are created through hydrostatic nucleosynthesis in the cores of massive stars. Long-lived isotopes (e.g. ^{26}Al , $t_{1/2} = 7.2 \times 10^5$ yr, $E_\gamma = 1809$ keV) can be convected to the surface and carried into the interstellar medium (ISM) by their stellar winds where they subsequently decay and produce gamma-ray lines. Such lines can be valuable tracers of recent nucleosynthesis in our galaxy.

Supernovae (SN): Radioactive nuclei are created by explosive nucleosynthesis in both Type I and Type II supernovae. In both cases the dominant element synthesized is ^{56}Ni which decays to ^{56}Co ($t_{1/2} = 6.8$ d, $E_\gamma = 158, 240, 270, 480, 811$ keV) and then to ^{56}Fe ($t_{1/2} = 79$ d, $E_\gamma = 847, 1238$ keV). Type Ia SN are believed to be produced by either the accretion-induced collapse of a white dwarf (WD) or the merger of a WD-WD binary. At early times the SN shell is optically thick to gamma-rays, becoming transparent after ~ 60 days. The strongest lines are therefore from the longer-lived decay of ^{56}Co . Later in the expansion even longer-lived isotopes take over (e.g. ^{57}Co , $t_{1/2} = 271$ d, $E_\gamma = 122, 136$ keV). Type II SN result from the core collapse of a more massive star which has exhausted its nuclear fuel. The envelope takes longer to become transparent and consequently the gamma-ray line strengths are generally weaker. As with Type Ia SN the $^{56}\text{Ni} \rightarrow ^{56}\text{Co} \rightarrow ^{56}\text{Fe}$ decay chain also dominates, followed by ^{57}Co decay. The former decay is believed to be the engine that powers the early-time expansion of the SN shell. ^{56}Co undergoes β^+ -decay 19% of the time. Most of the positrons will annihilate in the expanding SN shell before they can escape (Chan & Lingenfelter 1993; Milne et al. 1999). However, for Type Ia SN, if the magnetic field has a sufficiently smooth and open geometry, a significant number of positrons can escape. It is possible that these positrons can account for the observed bulge component of the Galactic 511 keV emission (Purcell et al. 1997; Knödlseeder et al. 2005).

Gamma-Ray Bursts/Hypernovae: Hypernovae are asymmetric explosions of massive stars. They are leading candidates for the origin of long duration gamma-ray bursts. It has been suggested that positron production by hypernovae can be a significant contributor to the Galactic 511 keV flux (Cassé et al. 2004; Bertone et al. 2004). Positrons can be efficiently transported outward by the strong jets that are believed to accompany hypernovae and escape into ISM where they subsequently annihilate.

Black Holes: Electron-positron jets are believed to be characteristic features of both super-massive and stellar-mass black holes (BH) (see, for example Rees (1984); Mirabel & Rodriguez (1999)). Positrons can be transported by the jets far from the BH where they eventually will annihilate in the ISM. The 511 keV bulge component could conceivably

originate from the massive BH at the center of our Galaxy or alternatively from an ensemble of stellar-mass BH's each producing positrons that migrate into the ISM and produce the diffuse emission that is observed. In the latter scenario, however, one might expect a more pronounced disk component with a larger disk-to-bulge ratio than is observed (Knödlseeder et al. 2005).

Pulsars: Electron-positron pairs are believed to be created by vacuum pair production in the strong magnetic fields in pulsar magnetospheres (see, for example, Ruderman (1972)). MeV positrons that escape the magnetosphere will take typically a few million years to annihilate. Depending on the geometry and smoothness of the local Galactic magnetic field the positrons could migrate far from the source before they annihilate. The resultant 511 keV emission could manifest itself as either an ensemble of point-like sources or a diffuse glow.

Neutron Star Binaries: Matter accreting onto a neutron star surface can produce gamma-ray lines by a variety of interactions. Probably the strongest of these is neutron-proton recombination which produces a line at 2223 keV (Bildsten, Salpeter & Wasserman 1993).

Cosmic Ray Interactions: Cosmic rays can activate nuclei in the ISM which subsequently decay and produce gamma-ray lines. Positrons can also be created through either β^+ -decay or π^+ production ($\pi^+ \rightarrow \mu^+ \rightarrow e^+$). The most extensive treatment of cosmic-ray produced gamma-ray lines can be found in Ramaty, Kozlovsky & Lingenfelter (1979). The authors have exhaustively inventoried the great number of different possible interactions and decay channels. The predicted gamma-ray spectra display a rich variety of lines in the energy range from a few hundred keV to 8 MeV. Narrow lines are normally produced by the interaction of cosmic-ray protons with ambient interstellar nuclei. Broadened lines are produced when cosmic-ray heavy nuclei collide with ambient interstellar hydrogen. Narrow lines are easier to detect with a high-resolution instrument such as INTEGRAL/SPI. The strongest narrow lines are those from electron-positron annihilation (511 keV), ^{56}Fe excitation (847 keV), ^{12}C excitation (4439 keV) and ^{16}O excitation (6129 keV). Unfortunately these line strengths are near the limits of detectability with current instrumentation. We will return to this point later with a comparison of our limits and the predictions of theory.

Dark Matter: There has been recent speculation on the possible existence of light ($m < 100$ MeV) dark matter particles which would decay or annihilate primarily through the formation of electrons and positrons (Boehm et al. 2004; Hooper & Wang 2004; Oaknin & Zhitnitsky 2004; Picciotto & Pospelov 2004). A variety of different particles including axions, sterile neutrinos and quark droplets have been proposed. Dark matter decay or annihilation could possibly account for the bulge component of the electron-positron anni-

hilation radiation. However, Beacom, Bell & Bertone (2004) have pointed out that dark matter annihilation or decay will be accompanied by gamma-ray production, which would exceed the observed gamma-ray flux from the Galactic Center for all but the lightest dark matter particles.

1.2. Known Astrophysical Gamma-Ray Lines

Table 1 summarizes the known astrophysical gamma-ray lines. The number of actually observed lines is small compared to the number of predictions. This may reflect an excess of optimism on the part of the theoreticians or insufficient sensitivity on the part of previous instruments. INTEGRAL/SPI represents a significant step forward in its ability to detect and study gamma-ray lines. Hence there is considerable motivation to thoroughly search it's data for new lines, which is the subject of this paper. Search variables include line energy, width, temporal and spatial distribution. INTEGRAL/SPI combines fine energy resolution with moderate spatial resolution. This, along with the need to examine different time scales (since most of the point sources are variable), leads to a very large volume of phase space to be searched. We elaborate the properties of the known lines as follows.

1.2.1. *Electron-Positron Annihilation*

The 511 keV electron-positron annihilation line from the GC was the first-discovered and remains the strongest. First detected in early balloon flights (Johnson, Harnden, & Haymes 1972; Leventhal, MacCallum, & Stang 1978), it has subsequently been studied in detail by several satellites [HEAO-C (Riegler et al. 1981), WIND/TGRS (Teegarden et al. 1996), CGRO/OSSE(Purcell et al. 1997)]. INTEGRAL (Jean et al. 2004; Teegarden et al. 2005; Knödlseeder et al. 2005) has given us the most recent and most detailed information. The results are all generally consistent with a narrow (~ 2 keV FWHM) unshifted line with the bulk of the emission coming from a $\sim 10^\circ$ diameter region centered at the Galactic Center. However, Jean et al. (2005) have recently reported evidence for a second broadened component of the 511 keV line such as would be expected from in-flight annihilation of positronium. OSSE and INTEGRAL also marginally detect a more extended Galactic Plane component (Purcell et al. 1997; Knödlseeder et al. 2005). In an early analysis OSSE (Purcell et al. 1997) reported an enhancement of 511 keV emission above the Galactic Plane (the so-called "annihilation fountain"). However, later OSSE analyses (Kinzer et al. 2001; Milne et al. 2001) found little or no evidence for this component, and INTEGRAL does not detect it (Jean et al. 2004; Knödlseeder et al. 2005). The strong concentration of emission and large

bulge-to-disk ratio imply that the Galactic 511 keV emission comes from an older Galactic population, for example SNIa or LMXBs (Knödlseeder et al. 2005).

In 1990-91 the SIGMA experiment on the Granat satellite found evidence for transient line-like features in the spectra of two different sources 1E1740.7-2942 (Bouchet et al. 1991) and Nova Musca (Goldwurm et al. 1992). In 1E1740.7-2942 a strong (flux = $1.2 \times 10^{-2} \text{ cm}^{-2} \text{ s}^{-1}$) broadened (FWHM = 240 keV) feature was found at an energy of 480 ± 90 keV. It lasted for ~ 2 days. It has been interpreted as (potentially) red-shifted $e^+ - e^-$ annihilation radiation. Four months later a similar event was observed from Nova Musca. The feature lasted \sim half a day. The measured line parameters were, flux = $6 \pm 3 \times 10^{-3} \text{ cm}^{-2} \text{ s}^{-1}$, $E = 481 \pm 22$ keV, and FWHM = 23 ± 23 keV. In subsequent years no further transient line-like features have been observed despite extensive coverage of the Galactic Center and Plane.

1.2.2. ^{26}Al

The 1809 keV line from radioactive ^{26}Al was first detected by HEAO-C (Mahoney et al. 1984). The COMPTEL experiment on CGRO produced detailed maps showing a distribution narrowly confined to the Galactic Plane with local hot spots in the Cygnus and Vela regions. The distribution was much more extended in the Galactic Plane than the 511-keV emission and was generally consistent with that expected for young massive stars (Diehl et al. 1995; Knödlseeder et al. 1999). A broadened line (6.4 ± 1.2 keV FWHM) was reported by the GRIS experiment (Naya et al. 1996); however subsequent observations by RHESSI (Smith 2004) and INTEGRAL (Diehl et al. 2003b) found no evidence for any line broadening. ^{26}Al undergoes β^+ -decay 82% of the time. The resultant 511 keV emission from the annihilation of these positrons is $\sim 50\%$ of the observed 511 keV disk component (Knödlseeder et al. 2005), but within the (large) errors consistent with it.

1.2.3. ^{44}Ti

^{44}Ti is synthesized in core-collapse supernovae. It undergoes the decay chain $^{44}\text{Ti} \rightarrow ^{44}\text{Sc} \rightarrow ^{44}\text{Ca}$ with a half-life of 59 yr and in the process produces gamma-ray lines at 68, 78 and 1156 keV, all of the same intensity. ^{44}Ti lines then are signatures of SN in our Galaxy occurring within the last several hundred years. COMPTEL (Iyudin et al. 1994) has detected the 1156 keV ^{44}Ti line from Cas A, the youngest known Type II SN remnant (age ~ 320 yr) in our Galaxy. Vink et al. (2001) reported the detection of the 68 and 78 keV

lines from Cas A using the BeppoSax satellite. The measured fluxes were consistent with that found previously by COMPTEL for the 1156 keV line. A subsequent detection of the 1156 keV line from the Vela region was reported by COMPTEL several years later (Iyudin et al. 1998). The location of the new COMPTEL source was found to be coincident with a young SN remnant discovered independently and nearly simultaneously by the ROSAT satellite (Aschenbach 1998).

1.2.4. ^{60}Fe

^{60}Fe is also believed to be produced in core-collapse supernovae. Its flux will accumulate from all of the Galactic supernovae occurring over its long half-life (1.5×10^6 yr). RHESSI (Smith 2004) and INTEGRAL/SPI (Harris et al. 2005) have both recently reported marginal ($3-4 \sigma$) detections of the 1172 and 1333 keV lines from ^{60}Fe . Theoretical models of Type II SN generally predict roughly the same amounts of ^{60}Fe and ^{26}Al (Limongi & Chieffi 2003). However, the $^{60}\text{Fe}/^{26}\text{Al}$ ratio reported by INTEGRAL/SPI is only 0.11 ± 0.03 (Harris et al. 2005), and RHESSI finds a similar value. The implication is that some other source of ^{26}Al besides explosive nucleosynthesis in core-collapse supernovae is required.

1.2.5. ^{56}Co , ^{57}Co

SN1987A, a type II SN in the Large Magellanic Cloud, was the nearest supernova in the era of modern astronomy when instrumentation existed capable of detecting gamma-ray emission. Gamma-ray lines from ^{56}Co decay ($t_{1/2} = 79$ d, $E_\gamma = 847, 1238, 2599$ keV) were first detected by the SMM satellite (Matz et al. 1988) ~ 6 months after the discovery of the SN. The competition between the increasing transparency of the SN shell and the decay of ^{56}Co produces an expected light curve for the gamma-rays with a maximum ~ 1 yr after the initial explosion. However gamma-ray emission was seen much earlier than expected, which argued for significant outward mixing of the ^{56}Co which was initially produced deep inside the supernova shell. Several high resolution germanium detectors were able to spectroscopically study the lines (Teegarden et al. 1988; Mahoney et al. 1988; Sandie, et al. 1988; Tueller et al. 1990). Although not entirely self-consistent, the ensemble of observations showed broadened (FWHM = 10-15 keV) and redshifted lines that were consistent with emission from an expanding shell of velocity ~ 3000 km/sec and substantial mixing.

Supernova models have suggested that ^{57}Co decay ($t_{1/2} = 271$ d) might be the dominant energy source driving the later stages of the shell expansion of a Type II SN (Clayton 1974;

Woosley, Pinto, & Hartmann 1989). In 1991 the OSSE experiment on the Compton Gamma-Ray Observatory (CGRO) reported the detection of the 122 and 136 keV ^{57}Co line complex from SN1987A (Kurfess et al. 1992). OSSE did not have sufficient resolution to resolve the lines from each other. The line flux implied that the ratio of $^{57}\text{Ni}/^{56}\text{Ni}$ produced in the supernova was ~ 1.5 times the solar system $^{57}\text{Fe}/^{56}\text{Fe}$ abundance ratio.

1.2.6. Solar Flares

Energetic particles accelerated in solar flares can excite nuclei in the solar atmosphere that subsequently decay emitting gamma-rays. These gamma-ray lines are useful diagnostics of the ambient conditions at the flare site and of the acceleration mechanism (for an extensive review see Ramaty & Murphy (1987)). The most extensive solar flare gamma-ray measurements have been carried by the GRS experiment on the Solar Maximum Mission (see, for example Forrest & Murphy (1988)). Gamma-ray lines from ~ 12 of the most abundant elements below and including Fe have been identified. Recently the RHESSI mission has provided high resolution spectroscopy using germanium detectors of solar-flare gamma-ray lines (Smith et al. 2003). RHESSI has succeeded in resolving profiles of gamma-ray lines from Fe, Mg, Ne, Si, C and O. The INTEGRAL instruments do not point at the sun, however for an intense solar flare there can be sufficient penetration of the shielding to allow useful gamma-ray line spectroscopy (Tatischeff et al. 2005)

2. Observations

2.1. The INTEGRAL Mission

INTEGRAL is a powerful space-borne observatory of the European Space Agency (ESA) designed to make detailed measurements of the spatial and spectral distribution of celestial hard X-rays and gamma-rays. For a detailed description of the mission see Winkler et al. (2003). Launched in mid-Oct. 2002, INTEGRAL has operated successfully up to the present day. The satellite was inserted into a highly eccentric orbit (10,000 km perigee, 160000 km apogee) with a period of 3 days. In this orbit INTEGRAL spends very little time inside the radiation belts and avoids the damaging proton belts entirely. INTEGRAL is operated as an observatory with $\sim 3/4$ of it's time open to guest observers. The remainder of the time (the so-called Core Program) is mainly dedicated to systematic deep exposures of the Galactic Center region and periodic scans of the Galactic Plane.

2.2. The INTEGRAL Spectrometer (SPI)

The Spectrometer for INTEGRAL (SPI) is one of the two major instruments on the INTEGRAL mission (Vedrenne et al. 2003). SPI employs an array of 19 high-resolution cooled germanium detectors in a close-packed hexagonal array. It covers the energy range 20 - 8000 keV, over which the energy resolution increases from ~ 2 to 8 keV FWHM. The detectors are cryogenically cooled by two dual Stirling cycle refrigerators to a nominal operating temperature of 86K. In addition to providing a precise determination of the photon energy, the Ge detector array also acts as an imaging device in conjunction with a tungsten coded-aperture mask located 1.7 m from the detector plane. Images are formed by deconvolving the shadow pattern cast by the mask on the detector plane. The fully-coded field-of-view of the telescope is 16° and the angular resolution is $\sim 3^\circ$ FWHM. The Ge detector array is surrounded by an active anti-coincidence shield/collimator made of bismuth germanate (BGO) scintillator. This dense hi-Z material is ideally suited for suppressing the strong flux of background radiation generated in both the SPI instrument and in the spacecraft. The active collimator defines a field-of-view of $\sim 25^\circ$ FWHM.

2.3. Data Selection

The data used in this investigation span the period from 2002 December to 2003 December. All INTEGRAL data becomes public after \sim one year so we have been able to use essentially all of the data from this period. The data have been carefully screened to remove bad periods due to enhanced solar activity, trapped radiation and noise. ~ 12 hr of data near perigee each 3 day orbit is not usable. In addition there were three SPI annealings during this period resulting in the loss of ~ 35 days of data. After all data selection criteria were applied the total remaining usable time in the first year was 12.3 Msec. A large fraction of the total observing time was spent in or near the Galactic Plane. 3.6 Msec were spent in the Galactic Center region (within 30° of the center) and 9.6 Msec were within $\pm 20^\circ$ of the Galactic Plane. An exposure map of the data used in this paper is shown in Fig. 1.

During nearly all observations the direction of pointing is stepped through a "dithering" pattern. This is to provide additional information that will be useful (essential in the case of SPI) for both determining the background and producing high quality images. For most point source observations a square 5×5 pattern with 2° step size was used. The duration of each pointing was typically ~ 2000 sec. For the Core Program deep exposure of the Galactic Center a raster scan was performed which typically covered $\pm 30^\circ$ in Galactic longitude and $\pm 20^\circ$ in Galactic latitude with step size and duration similar to that for point sources. Our total data set comprises 6194 of these pointings.

2.4. Instrument Background

Background is primarily produced in SPI through the interaction of cosmic rays with the telescope and with surrounding material in the spacecraft (for detailed discussions see Jean et al. (2003); Weidenspointner et al. (2003); Teegarden et al. (2004)). The diffuse cosmic background also contributes mainly at energies below ~ 200 keV. Cosmic rays colliding with nuclei in the detectors and surrounding material can fragment them into many different kinds of isotopes. Many of these are created in excited unstable states which subsequently decay with the emission of gamma-rays (and other particles as well). Gamma-rays are produced at (nearly) discrete energies corresponding to the transitions between various nuclear levels. This leads to the complex multi-line SPI background spectrum shown in Fig. 2. For a nearly complete identification of the SPI background lines see Weidenspointner et al. (2003). The strongest background lines as well as certain lines that fall at astrophysically interesting energies (e.g. 511 keV, 1809 keV) are labeled with their energies in keV. Below ~ 200 keV the spectrum is dominated by very strong lines (54, 67, 139, 198 keV) that are primarily long-lived isomeric states which are not rejected by the SPI active anti-coincidence shield. These lines account for $\sim 30\%$ of the total SPI background counting rate. There is a strong background line at 511 keV which is particularly problematic for the astrophysical 511 keV analysis. A weaker background line is present at 1809 keV which must also be taken into account. Above 6 MeV a number of lines are seen corresponding to the highest energy levels of the isotopes of Ge. Here also, one sees the effects of instrumental line broadening which increases monotonically with energy. The complex feature between ~ 1400 and 1600 keV is an instrument artifact most probably arising from very large (> 1 GeV) energy losses in the Ge detectors due to nuclear spallation and/or heavy cosmic rays. We have excluded this region from our analysis.

The cosmic rays producing the SPI background vary in intensity due to modulation by the solar magnetic field. The SPI background responds in a complex way to these variations. In prior analyses (Jean et al. 2003; Teegarden et al. 2004) the rate of events saturating the Ge detectors (GEDSAT) was found to be a good tracer of the SPI background. It is dominated by energetic particles (i.e. cosmic rays) with $E > 8$ MeV. This rate is plotted in Fig. 3. During this one year period the total variation is $\sim 10\%$. Over longer time intervals larger variations are expected due to the 11 yr period of the solar cycle. Most of the variability is on time scales longer than a few days. We will return to this point later as it is crucially important to our methods of background subtraction.

SPI is a sensitive instrument and with long accumulation times is capable, in principle, of detecting very weak gamma-ray line fluxes. However, at its limiting sensitivity (for, say, $T = 10^6$ sec), the typical signal-to-background ratio is only a fraction of a percent. Using a

tracer such as the GEDSAT rate we can normalize and subtract the background. However, the subtraction is not perfect. In the simplest case the intensity of a particular background line is determined by a convolution of the cosmic-ray spectrum with the energy-dependent cross section for the production of that line. The shape of the cosmic-ray spectrum changes with changing solar modulation, which means that there will not be a simple linear dependence between the cosmic-ray and instrument background rates. In addition a significant number of background-producing isotopes have long (> 1 day) half-lives. This again means that the SPI background will not simply track the cosmic-ray intensity. Our approach will be to choose the background dataset near in time ($\lesssim 20$ days) to the source observation. We will see later that this can very effectively suppress the background-subtraction residuals. It will be necessary to trade off suppression of residuals against loss of sensitivity due to the reduced availability of background data. We also found that the standard procedure for gain correction, i.e. performed only once per orbit, was not adequate. There were times when gain variations within an orbit could produce significant residuals in the background-subtracted spectra. We therefore performed a secondary gain correction on a pointing-by-pointing basis which improved the background subtraction significantly.

3. Data Analysis

3.1. General Search Procedure

The line search procedures used in this study varied according to whether or not the sources were diffuse or point-like, steady or variable. For diffuse emission they also varied with the scale of the assumed source. However, the following elements were common to all procedures. First, a background-subtracted spectrum was created. This is the most crucial step. As was seen in Fig. 2 there are numerous strong background lines. These lines vary in strength and relative intensity. Also, the line profiles change with time due to radiation damage. All of these effects can produce residuals in the background-subtracted spectra that can masquerade as astrophysical lines. A number of different background subtraction methods were investigated. The most effective method and final choice was the application of time-dependent background selection criterion as discussed in the preceding section.

The background-subtracted spectra were all then run through the same search procedure. The spectra were convolved with the template shown in Fig. 4. The effect of this convolution is to subtract the local continuum from the line, which is necessary when there is a significant continuum from the source. It also removes any residual continuum due to imperfect background subtraction. The convolution was repeated while varying the width ΔE of the convolution kernel over the range 3-1000 keV in 20% increments. Various quality

checks were included the software. For example, bipolar residuals, which can be produced when the width of a background line varies, were rejected. As an aid, software tools were created that allowed the rapid viewing of large numbers of spectra to verify the procedures and examine candidate lines.

3.2. Diffuse Sources

For the diffuse source search we use the so-called "light bucket" approach. In this method all 19 SPI detectors were summed, effectively creating a simple collimated instrument whose field-of-view ($\sim 25^\circ$ FWHM) is determined by the SPI active BGO collimator. A summed raw spectrum was created for each of the 6194 pointings. From these spectra we created two different kinds of background-subtracted spectra.

3.2.1. Large Scale ($\gtrsim 20^\circ$)

By large-scale we mean spectra accumulated over large scale regions, specifically the Galactic Center (within 13° or 30° of the center) and Galactic Plane (within $\pm 20^\circ$ of the plane excluding the central region). The available source pointings were accumulated into daily spectra and background pointings for each source spectrum were chosen according to the following criteria: 1) the pointing is not in the source region and 2) the time interval between source and background pointings was < 20 days. The choice of 20 days was based on the trade-off discussed earlier between residuals in the spectrum and loss of sensitivity due to the reduction in the available background data. Background spectra were normalized to source spectra before subtraction using the 198 keV instrumental line, which provided the best cancellation of the strong background lines. Finally an average was taken over all of the Galactic Center daily spectra and separately over the Galactic Plane daily spectra. The daily spectra may have common background pointings and are thus not necessarily independent. This was taken into account in the error and sensitivity calculations.

Fig. 5 shows large-scale difference spectra formed by subtracting all the off-center data ($> 30^\circ$ from the center) from the on-center data. The 20-300 region containing the strongest background lines is plotted. The individual data points are the background-subtracted count rates. In the upper panel the background (off-center data) has been selected as described in the preceding paragraph using a value of 20 days for the maximum allowed time difference between source and background. In the lower panel no background selection has been applied and strong residuals are evident. The 198 keV line has been used for normalization so that

its net residual is zero. However, there is a significant bipolar feature due to the changing shape of the line (a radiation damage effect). Other bipolar residuals are seen at 93 keV and 139 keV. The strongest residual is from the 143 keV background line (a rather weak one). The reason for it's prominence is that it is produced by the decay of ^{57}Co with a half-life of 271 days. The strength of this line will build up over this period and it will not track the variations of the other lines. This spectrum illustrates the critical importance of proper selection of background data. Also shown in the upper panel is the raw unsubtracted spectrum multiplied by 0.01. The effectiveness of the line cancellation is evident - the residuals are typically $\lesssim 0.1\%$.

3.2.2. *Small-to-Medium Scale ($\lesssim 20^\circ$)*

To search for smaller scale diffuse emission a series of uniformly-spaced grid points on 18° centers in Galactic coordinates was defined. For each point on the grid source data were taken to be all pointings with 12° (the half-response) of that point. As before, the available source pointings were binned in 1 day intervals thus creating a set of daily source spectra for each grid point. Background spectra for each grid point source spectrum were chosen as follows: 1) the angular distance from the grid point was $> 20^\circ$ and 2) the time interval between source and background pointings was < 3 days. In this analysis it was possible to use a smaller 3 day limit since background pointings were generally more readily available than they were for the large-scale analysis. Also, as before, an accumulated average spectrum was calculated for each grid point. Both the individual daily spectra and the average spectra were searched for lines following the procedure of Section 3.1. Although intended primarily for the search for diffuse lines, this procedure could also detect lines from point sources in the neighborhood of each grid point if they were present. It could conceivably catch point sources missed in the SPIROS-based analysis described in Section 3.3.

3.3. Point Sources

The SPI Bright Source Catalog¹ containing 159 sources was used as a primary input for the search for gamma-ray lines from point sources. The SPI standard analysis pipeline software was used [OSA Version 5.0, Diehl et al. (2003a); Walter et al. (2004)]. Point source spectra were created by the SPIROS program [SPI Iterative Removal of Sources (Dubath et al. 2005)]. The program initially creates a correlation map. The brightest source is

¹http://heasarc.gsfc.nasa.gov/docs/integral/bright_sources.html

identified, convolved through the response function and subtracted from the data. This process is repeated until no further significant sources remain. In this manner the effects of ghost images from bright sources in the field-of-view are suppressed or eliminated. The spectral binning (1 keV, 20-2000 keV; 10 keV, 2000-8000 keV) has been carefully chosen to ensure that there are sufficient counts per bin to preserve gaussian statistics (and therefore a linear fitting procedure) while maintaining the finest possible resolution to allow searching for the narrowest possible lines. In addition to the per orbit spectra a cumulative average spectrum over all the available data was determined. Finally, we also have created stacked spectra over generic types of sources (LMXB, HMXB, pulsar, SNR, Seyfert). In the per-orbit spectra the background is fit at the same time as the source, which means that it comes from the same orbit as the source. The background subtraction is therefore immune to most of the long-term effects described in Section 2.4.

In Fig. 6 we show spectra from the Crab and Cyg X-1. These are the principal INTEGRAL calibration targets and were observed for long periods of time. We have chosen these long accumulations since they best illustrate the effects of residuals due to the imperfect removal of instrumental background lines. Again the spectra are plotted with 1 keV binning in the 20-300 keV interval, which is the region where the strongest background lines are located. For the Crab spectrum with an integration time of 3.17 ksec there is a line-like residual at 198 keV, the energy of the strongest background line. The error bars are large so that the significance of this feature is not great. No other significant residuals appear in the Crab spectrum. For Cyg X-1 with a longer integration time (1.19 Msec) more residuals appear at 54, 67, 139 and 198 keV. These spectra illustrate the limiting factor the line search. Spectra of durations up to a few hundred ksec are relatively systematics free. For spectra of longer duration the number of line-like artifacts increases. We can (and have) searched these spectra, but we needed to perform an additional screening to remove the residuals at the locations of the strongest background lines.

4. Results

4.1. Diffuse Sources

4.1.1. Large Scale Diffuse Emission

We created three different large-scale background-subtracted spectra with parameters summarized in Table 2. Our standard line search algorithm was applied (see Section 3.1) to these three spectra with a 4σ threshold. Several line-like residuals were found at the positions of strong background lines, but no new lines appeared at other energies. The 511 keV

and 1809 keV lines were detected at the expected fluxes and significances. For large-scale diffuse emission the flux sensitivities depend on the assumed shape of the spatial distribution. We calculated sensitivities for four different distributions summarized in Table 3. The 10° gaussian (I) was chosen since it characterizes the known distribution of the $e^+ - e^-$ annihilation radiation. The flat circular distribution (II) was included as an unbiased sample of the central radian of the galaxy. The flat longitudinal distribution (III) is used as unbiased sample of the Galactic Plane. The Dirbe $60\mu\text{m}$ distribution was included as representative of the massive star distribution in our Galaxy since massive stars are a likely source of nucleosynthetic radioactive nuclei. It is one of the best matches to the observed ^{26}Al 1809 keV spatial distribution (Knödseder et al. 1999). Finally, the Egret 30-100 MeV distribution was used as representative of the cosmic-ray + interstellar gas distribution. In Fig. 7 we show the sensitivity as a function of energy for the Galactic center data set (13° radius) and 10 deg gaussian case (GC13,I). The curves are parametrized by line width (FWHM). The large-scale sensitivity results are summarized for all the appropriate combinations of data sets and flux distributions in Table 4. Narrow-line sensitivities are given at several selected energies. Also given are scale factors by which to multiply the curves of Fig. 7 in order to obtain sensitivities at other energies and line widths. and the scale factor to use in determining the sensitivities at other energies and line widths.

4.1.2. *Small-to-Medium Scale Diffuse Emission*

Daily background-subtracted spectra were created for each Galactic grid point as discussed in Section 3.2.2. We tested the spectra for well-behaved errors by comparing the standard deviations the data points with the predicted statistical errors. This was done in several energy bands over the entire spectral range. Fig. 8 shows the energy-averaged ratio of standard deviation to statistical error for each Galactic grid point plotted as a function of the total accumulation time. For purely statistical errors this ratio should be unity independent of accumulation time. For accumulation times $\gtrsim 10^5$ sec a moderate departure from unity is seen. This is interpreted as due to the decreasing statistical error becoming comparable to the systematic error due to imperfect background subtraction. To deal with this we have chosen a threshold of 5σ for line detection in the "light bucket" mode.

Fig. 9 is an example of the sort of plot we used to scan the data for interesting line candidates. The vertical axis is the line energy and in this case the horizontal axis is the Galactic longitude of the grid point. (Here the Galactic latitude is suppressed to allow creation of a 2D plot.) The vertical bars seen in some points are the line widths (not errors). The symbol size reflects the significance of the line detection. The 511 keV and 1809

keV lines are identified. They are detected at the correct location, energy, intensity and significance, which validates our search procedure. The 511 keV line is the most significant of all those appearing in the plot. Two lines (535 keV and 7430 keV) are detected at > 6 sigma. However, both of them coincide with strong instrumental lines. A number of other line candidates are seen, however they are all near the 5σ threshold. We have individually examined all of these lines. We find no lines at astrophysically interesting energies, nor do we find any clustering of candidates in our energy-longitude space. We therefore interpret all of these as spill-over from our 5σ threshold due to imperfect background subtraction.

In Table 5 we summarize the SPI sensitivities at each of our Galactic grid points. We show only those points for which there was a significant exposure ($T > 10^4$ sec). As in Table 4 we give the narrow-line sensitivities at several selected energies including 511 keV and the scale factor to use in determining the sensitivities at other energies and line widths.

4.2. Point Sources

Fig. 10 shows the line sensitivity for the crab as a function of energy for four different line widths. This plot is to be used as a template (similar to Fig. 7) for determining the sensitivities of other point sources. A total of 2661 spectra from 143 different sources were created and searched. Tables 6 to 12 summarize the results. The tables are organized by type of source (LMXB, HMXB, Pulsar, Quasar, SNR). Where it was relevant an average spectrum over all sources within the given type was created and searched for lines. The Seyferts (Table 11) span a range in redshifts of 0.001 - 0.063. When the average Seyfert spectrum was created a redshift correction was applied. As before, narrow (< 3 keV) line sensitivities are given for four selected energies and a scale factor (to multiply Fig. 10) is included for determination of sensitivities at other line energies and widths. The tables show only the results for the total accumulated spectra for each source. We also created and searched individual spectra for each revolution (3 day period) in which the source was observed. Finally we selected observations for each source in which the source was detected by SPI ($> 6\sigma$) in the 20-50 keV range and created and searched spectra from this subset. The results are too numerous to include here. Suffice it to say that no lines were detected in any of these spectra.

4.3. Rapidly Variable Sources

The line search discussed in Section 4.1.2 was sensitive to sources which vary on time scales $\gtrsim 1$ day. We also conducted a search for sources that vary on shorter time scales. A convenient time interval is that corresponding to a single pointing (typically 1000-2000 sec duration). The data and standard software are organized in a manner that makes such a search relatively easy. Spectra were created for each pointing and screened as before for solar and magnetospheric particle contamination. A total of 6004 spectra were produced. Background was simply taken to be the 1-day average spectrum from the day in which the pointing occurred. To maintain well behaved errors, we found it necessary to rebin the > 2000 keV part of the spectrum into 10 keV bins. We are therefore less sensitive to lines of width < 10 keV having energy > 2000 keV, but the > 2000 keV instrumental resolution is not too much smaller than 10 keV, so not much is lost. The same line search algorithm described in Section 3.1 was used to search all of the spectra. Fig. 11 shows the results of the search. It is similar to Fig. 9 except that the horizontal axis is time (IJD = days from 1/1/2000) rather than Galactic longitude. Here we have used a 4σ threshold rather than 5σ as before since with these shorter accumulations the effects of systematic errors due to imperfect background subtraction are smaller. We have individually examined all spectra in which line candidates are found. The most frequent detections are lines of ~ 60 keV energy and $\sim 15 - 20$ keV width. These are clearly residuals from the strong background line complex in the 54-67 keV range (see Fig. 2). The strongest detections are a pair of lines at 58 and 76 keV. This pair appears 5-7 times in the plot. After examining the data we found that they are produced by energetic electrons that leaked from the earth's magnetosphere and escaped our filter. The detections are fluorescence lines activated by bremsstrahlung photons from the magnetospheric electrons. The 58 keV line is the K_{α} line from tungsten (the material of the mask) and the 76 keV line is K_{α} line from bismuth (one of the main components of the shield). Again, these detections serve to validate our procedures. All other detections are just above the 4σ threshold and are expected since the number of trials is large (6004 spectra x 2800 energy bins/spectrum). From this analysis and that described in Section 4.1.2 we conclude that there is no evidence for variable lines in the INTEGRAL/SPI data on time scales from a fraction of an hour to several months.

5. Discussion and Conclusions

We have carried out a comprehensive search of the first year of the INTEGRAL/SPI data for gamma-ray lines and found no new lines. Our search has spanned a broad range in line energy ($\sim 40 - 8000$ keV), width (3 - 1000 keV) and variability time scale (~ 1000

sec. to constant). We have searched for both point-like and diffuse sources. We used the INTEGRAL Bright Source Catalog as an input and searched the spectra of 143 known sources. The diffuse source search covered the spatial range from $\sim 20^\circ$ (the approximate field-of-view of the spectrometer) up to the entire Galactic plane. This search would also have detected previously unknown point sources with line fluxes exceeding our detection threshold. The detection thresholds varied over a wide range (a few $\times 10^{-5}$ to a few $\times 10^{-3}$ $\text{cm}^{-2} \text{s}^{-1}$) depending on line energy and width, spatial scale, and variability time scale.

We discussed various predictions of gamma-ray lines in Section 1.1. In most cases the predictions were below the current INTEGRAL/SPI sensitivity thresholds. For pulsars, Romani (1996) has predicted a positron production rate of $\sim 10^{38} \text{s}^{-1}$, which leads to a 511 keV flux of $\sim 2 \times 10^{-7} \text{cm}^{-2} \text{s}^{-1}$ at the earth. For comparison, the INTEGRAL/SPI 511 keV sensitivities for two pulsars PSR B1509-38 and Vela (see Table 8) are $1 - 2 \times 10^{-4} \text{cm}^{-2} \text{s}^{-1}$, well above the prediction. For neutron stars, Bildsten, Salpeter & Wasserman (1993) have investigated the production of gamma-rays through accretion-induced surface reactions. The strongest predicted line is that from neutron-proton recombination (2223 keV). Neutrons are created by spallation of accreting helium and heavier nuclei. The predicted flux from Sco X-1 is $\sim 10^{-6} \text{cm}^{-2} \text{s}^{-1}$ with fluxes as high as $\sim 10^{-5} \text{cm}^{-2} \text{s}^{-1}$ being possible. The INTEGRAL/SPI 2.2 MeV sensitivity for Sco X-1 can be found by applying the scale factor in Table 6 to Fig. 10. The derived value is $\sim 10^{-4} \text{cm}^{-2} \text{s}^{-1}$, about an order of magnitude larger than the most optimistic prediction. The INTEGRAL Sco X-1 sensitivity can be expected to improve over the course of the mission since we have only analyzed the first year of data in which there were $\sim 7 \times 10^5$ sec. of useful data from Sco X-1.

Perhaps the most promising scenario for the detection of new gamma-ray lines is from cosmic-ray excitation of ambient nuclei in the interstellar medium. This process has been treated extensively by Ramaty, Kozlovsky & Lingenfelter (1979). The principal uncertainties are the composition of the cosmic rays and the ambient medium and the energy spectrum and density of the cosmic rays. The strongest line predicted by Ramaty, Kozlovsky & Lingenfelter (1979) is from the de-excitation of ^{12}C (4439 keV). The largest predicted flux is $\sim 7 \times 10^{-5} \text{cm}^{-2} \text{s}^{-1} \text{rad}^{-1}$ from the central region of our Galaxy. The most directly comparable line search result is that in Table 4 for the GC30, II case where the flux is assumed to be uniformly distributed over a circular region at the Galactic Center of radius 30° . The INTEGRAL/SPI sensitivity at 4439 keV is $\sim 10^{-4} \text{cm}^{-2} \text{s}^{-1}$ and the value is roughly the same if calculated per radian. This value is only slightly higher than the most optimistic prediction of Ramaty, Kozlovsky & Lingenfelter (1979). Over the course of the mission INTEGRAL will spend a sizable fraction of its time observing the central region of our Galaxy, and we can expect that the limits will significantly improve and provide a more stringent test of the model.

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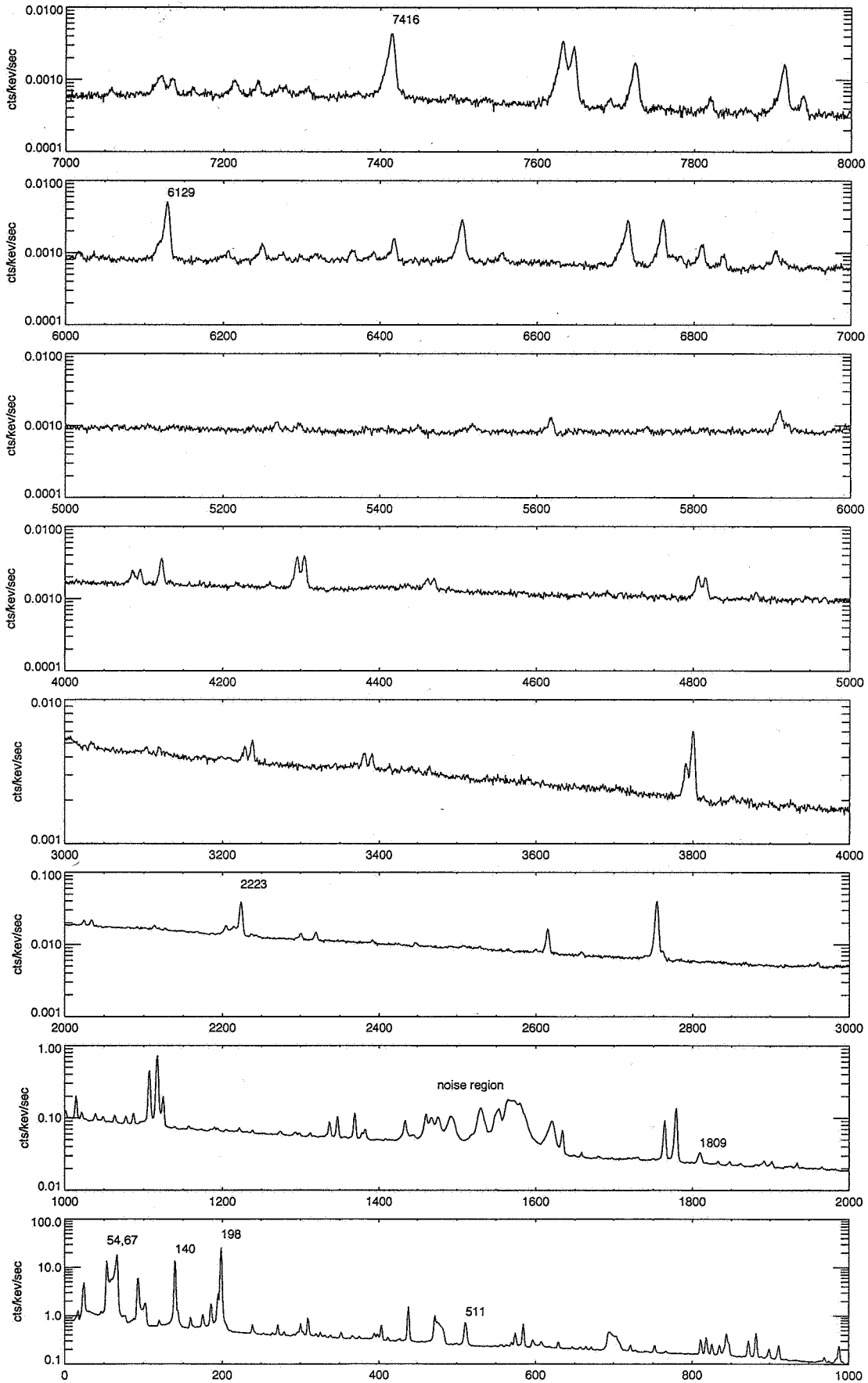
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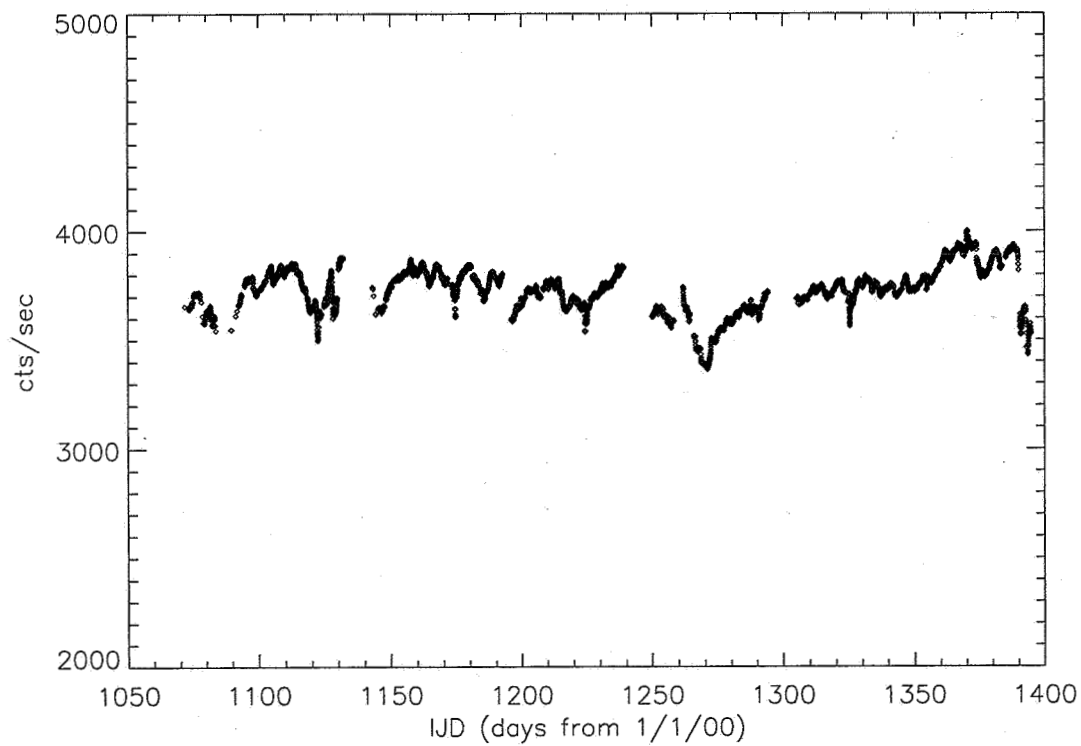


Fig. 3.— INTEGRAL/SPI germanium detector saturating event rate. This shows the typical temporal behavior of the SPI background.

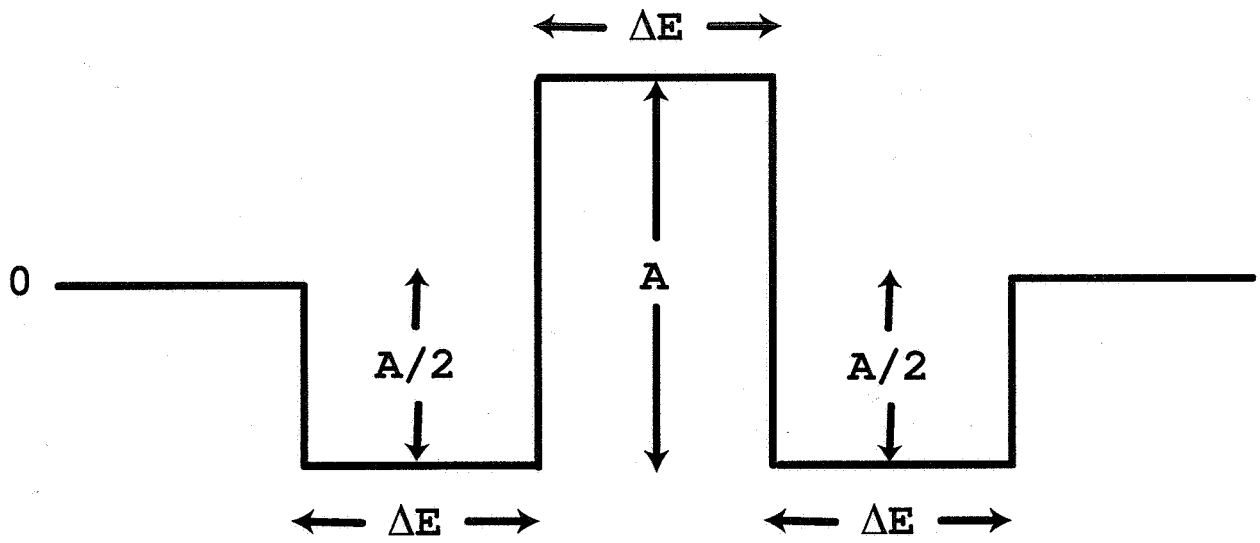


Fig. 4.— Convolution template used in line search.

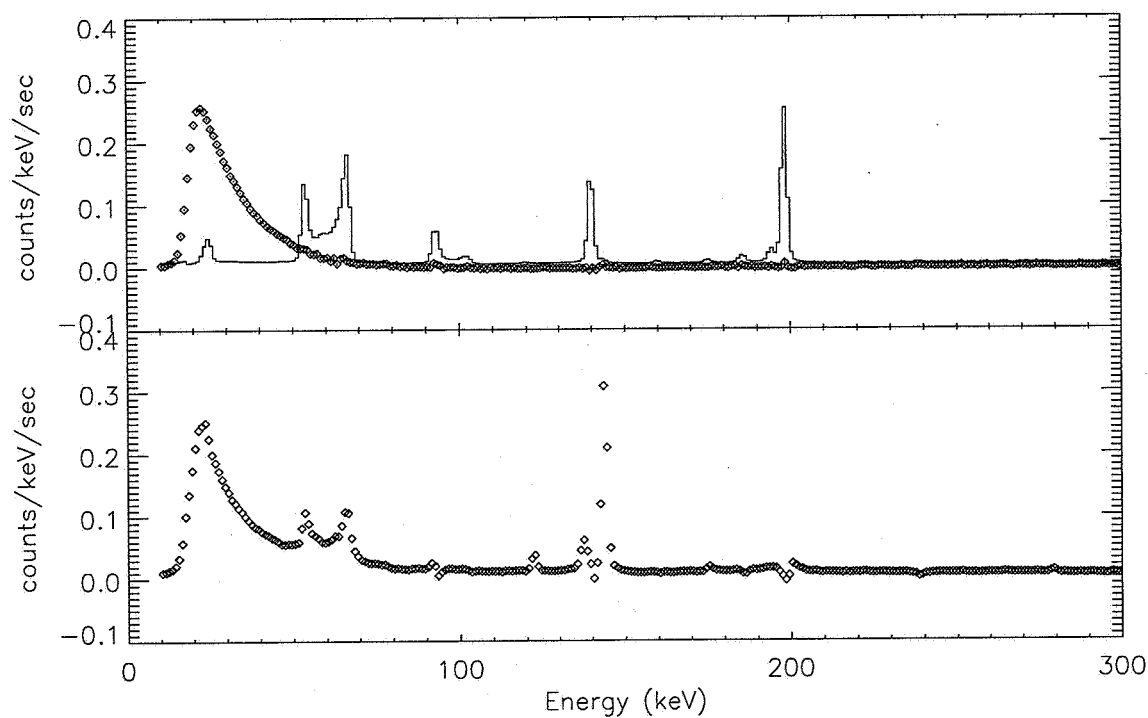


Fig. 5.— Galactic Center - Off-Center difference spectra. Upper panel: 20 day background selection criterion. Lower panel: no background selection criterion. Dramatic suppression of background line residuals is evident with 20 day selection criterion. Solid line in upper panel is 1% of total background spectrum. Residuals at the locations of the strong lines are $\lesssim 0.1\%$ of the raw background.

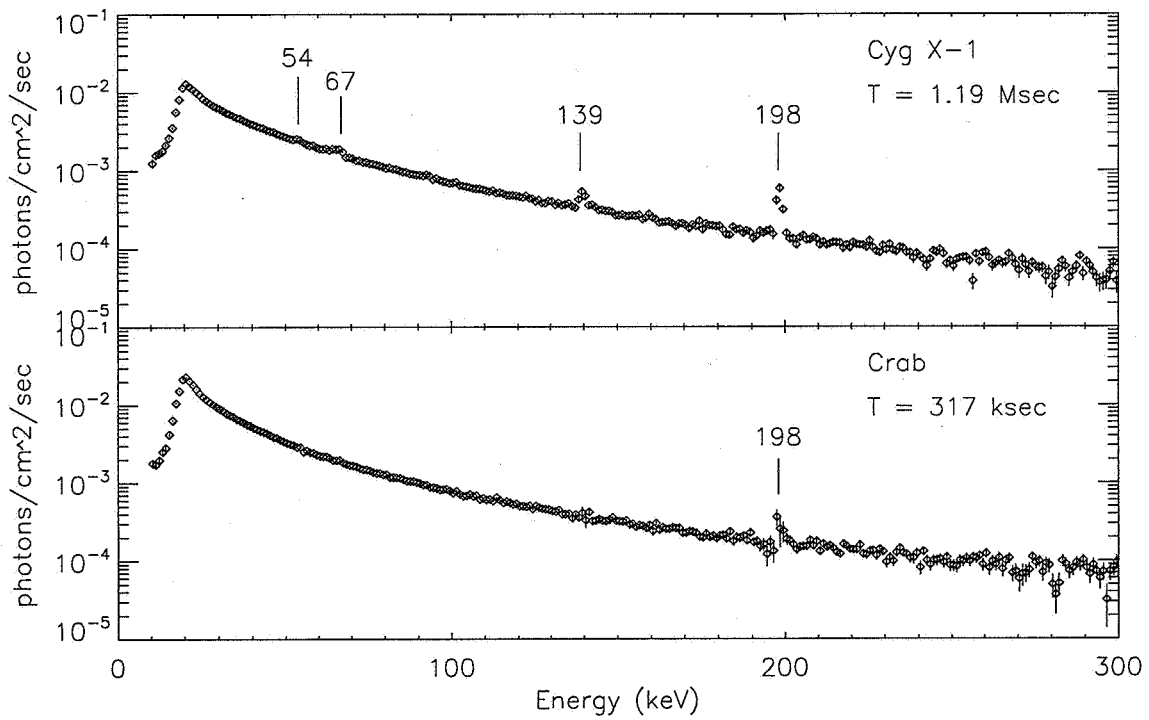


Fig. 6.— SPIROS spectra of the Crab and Cyg X-1 in the 20-300 keV range with 1 keV binning. T = spectrum integration time. Labelled lines are residuals of strong instrumental background lines.

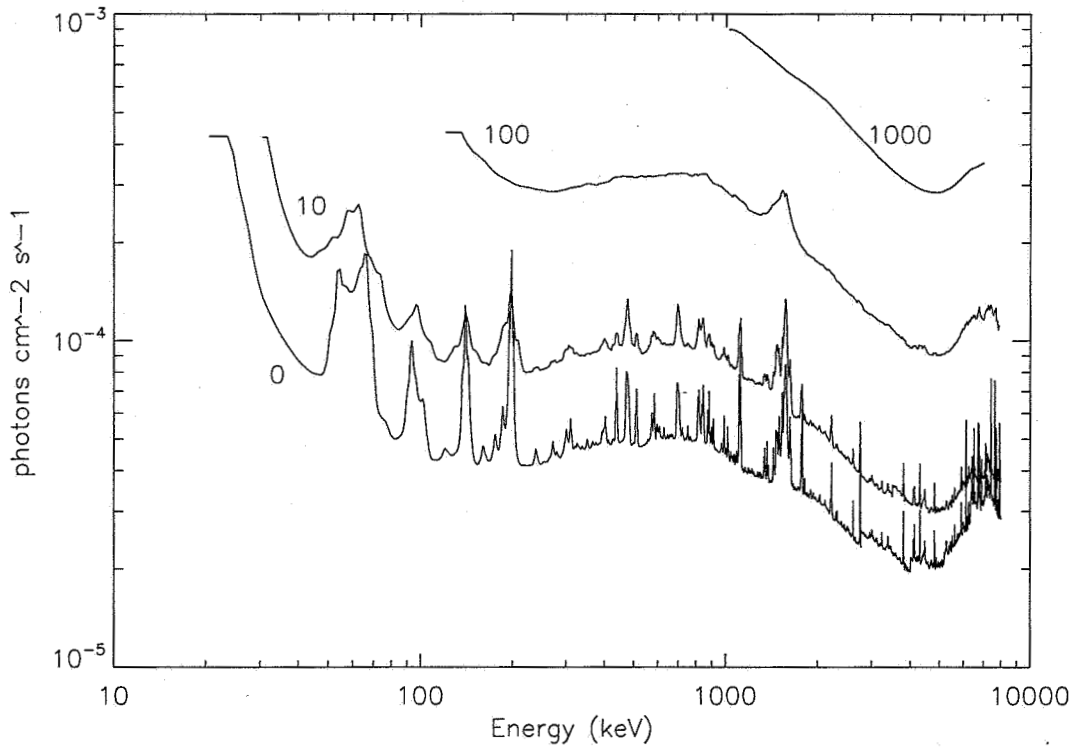


Fig. 7.— Line search sensitivity (3.5σ) for Galactic Center spectral accumulation ($\theta < 13^\circ$) and assumed 10° gaussian spatial distribution (GC13,I). Curves are labeled with assumed line width (keV FWHM).

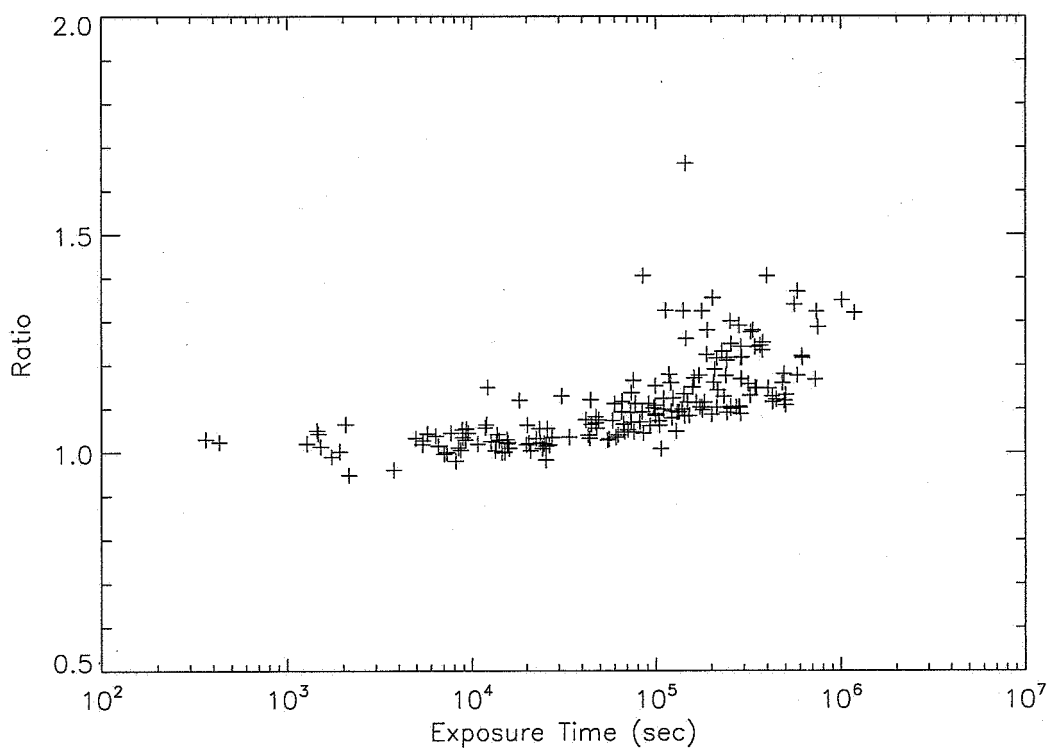


Fig. 8.— Energy-averaged ratio of standard deviation to median error of data points of difference spectra (similar to Fig. 5) vs. spectrum accumulation time. Spectra with long accumulation times display larger ratios reflecting increasing importance of systematic residuals.

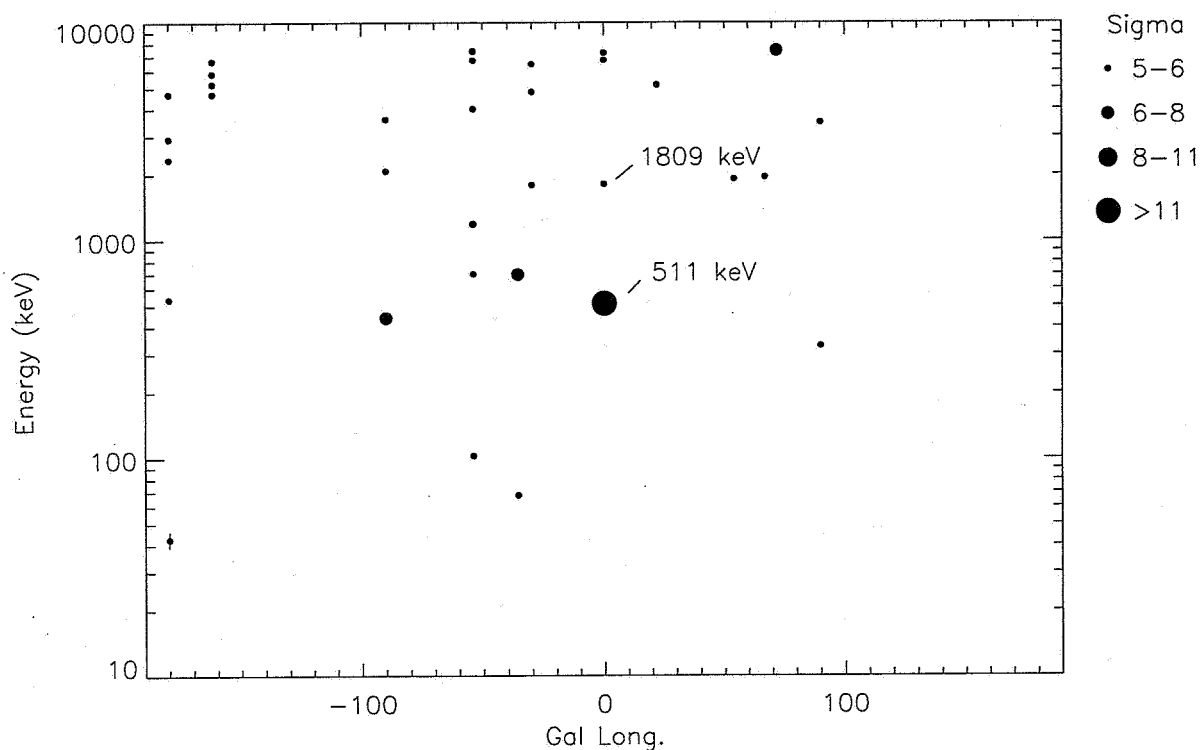


Fig. 9.— Results of small-to-medium scale diffuse line search using program described in Section 3.1. These results are derived from the accumulated average of the daily spectra for each Galactic grid point. Horizontal axis is Galactic longitude. Galactic latitude is suppressed in order to allow creation of 2D plot. Vertical axis is line energy. Vertical bars through data points show line width (not error). The ^{26}Al (1809 keV) and $e^+ - e^-$ annihilation (511 keV) lines are labelled. These lines are detected by our program at the correct location, energy and intensity.

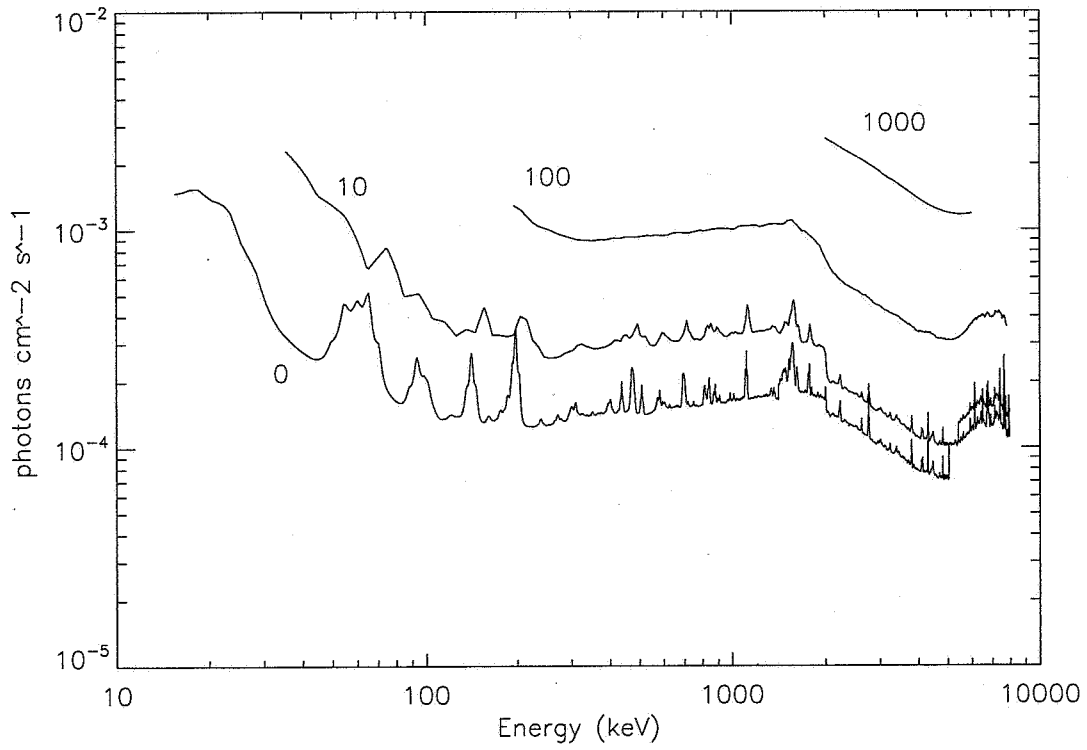


Fig. 10.— Line search sensitivity (3.5σ) for Crab. Curves are labeled with assumed line width (keV FWHM).

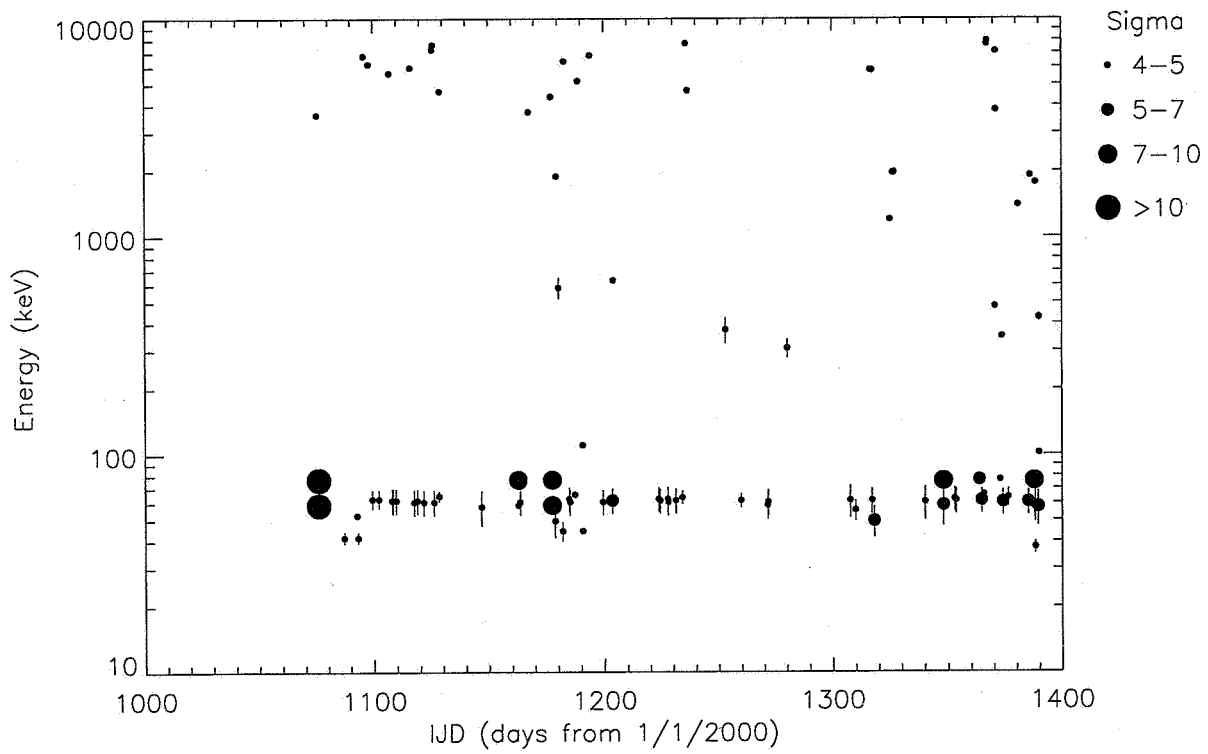


Fig. 11.— Results of rapidly variable line search using program described in Section 3.1. Vertical axis is line energy. Vertical bars through data points show line width (not error). The strong line pairs at 58 and 76 keV are from occasional bursts of magnetospheric electrons.

Table 1. Observed Astrophysical Gamma-Ray Lines

Line	Energy (keV)	Half-Life (yr)	Source(s)
$e^+ - e^-$ annihilation	511	N/A	Galactic Center/Plane 1E1740.7-2942 Nova Musca
^{26}Al	1809	7.2×10^5	Galactic Plane
^{44}Ti	68, 78, 1156	63	Cas A Vela Region
^{60}Fe	59, 1172, 1333	1.52×10^6	Galactic Plane
^{56}Co	847, 1238	0.21	SN1987A
^{57}Co	122, 137	0.74	SN1987A
^2H , ^{12}C , ^{16}O , ...	2223, 4438, 6129, ...	prompt	Sun (solar flare)

Table 2. Large-Scale Data Sets

Label	Description	Criterion
GC13	(Galactic Center) - (Off-Center)	GC region: $< 13^\circ$ from center
GC30	(Galactic Center) - (Off-Center)	GC region: $< 30^\circ$ from center
GP	(Galactic Plane) - (Off-Plane)	GP region: $\ b\ < 20^\circ$

Table 3. Diffuse Source Spatial Distributions

Label	Description
I	symmetric gaussian (10° FWHM) centered at GC
II	flat distribution (radius = 30°) centered at GC
III	flat in Galactic longitude, gaussian in Galactic latitude (FWHM = 5°)
IV	Dirbe $60\mu\text{m}$
V	Egret 30 - 100 MeV

Table 4. Large-Scale Sensitivities

Data Set	Spatial Distribution	Exposure (sec)	100 keV ^a	500 keV ^a	1000 keV ^a	5000 keV ^a	Scale Factor ^b
GC13	I	1.905×10^6	6.53×10^{-5}	4.87×10^{-5}	4.61×10^{-5}	2.07×10^{-5}	1.00
GC30	II	3.621×10^6	2.22×10^{-4}	1.50×10^{-4}	1.30×10^{-4}	4.80×10^{-5}	3.08
GP	III	3.846×10^6	6.22×10^{-4}	4.41×10^{-4}	4.01×10^{-4}	1.64×10^{-4}	9.05
GP	IV	3.846×10^6	3.17×10^{-4}	2.26×10^{-4}	2.05×10^{-4}	8.37×10^{-5}	4.63
GP	V	3.846×10^6	9.35×10^{-4}	6.43×10^{-4}	5.70×10^{-4}	2.20×10^{-4}	13.21

^aNarrow line sensitivity [total Galactic flux (photons $\text{cm}^{-2} \text{sec}^{-1}$)]

^bScale factor to apply to sensitivity plot of Fig. 7.

Table 5. Small-to-Medium-Scale Sensitivities

l	b	Exposure (sec)	100 keV ^a	511 keV ^a	1000 keV ^a	5000 keV ^a	Scale Factor ^b
180.0	-54.0	2.560×10^5	1.86×10^{-4}	1.85×10^{-4}	1.37×10^{-4}	6.30×10^{-5}	2.16
-180.0	-18.0	1.309×10^5	3.53×10^{-4}	3.45×10^{-4}	2.54×10^{-4}	1.17×10^{-4}	4.02
180.0	0.0	4.191×10^5	1.69×10^{-4}	1.66×10^{-4}	1.21×10^{-4}	5.56×10^{-5}	1.93
-180.0	72.0	3.656×10^4	5.44×10^{-4}	5.45×10^{-4}	3.99×10^{-4}	1.83×10^{-4}	6.36
-162.0	0.0	3.666×10^5	1.80×10^{-4}	1.77×10^{-4}	1.30×10^{-4}	5.93×10^{-5}	2.07
-158.0	-18.0	4.880×10^4	4.39×10^{-4}	4.30×10^{-4}	3.17×10^{-4}	1.46×10^{-4}	5.02
-144.0	0.0	2.250×10^4	6.19×10^{-4}	6.14×10^{-4}	4.53×10^{-4}	2.09×10^{-4}	7.17
-126.0	0.0	2.507×10^4	5.79×10^{-4}	5.77×10^{-4}	4.26×10^{-4}	1.97×10^{-4}	6.74
-108.0	0.0	6.608×10^4	3.78×10^{-4}	3.77×10^{-4}	2.81×10^{-4}	1.30×10^{-4}	4.40
-90.0	-36.0	3.837×10^5	1.50×10^{-4}	1.49×10^{-4}	1.11×10^{-4}	5.11×10^{-5}	1.74
-90.0	0.0	1.439×10^5	2.40×10^{-4}	2.40×10^{-4}	1.79×10^{-4}	8.29×10^{-5}	2.80
-90.0	54.0	1.281×10^5	2.58×10^{-4}	2.58×10^{-4}	1.92×10^{-4}	8.87×10^{-5}	3.02
-90.0	72.0	1.933×10^5	2.11×10^{-4}	2.10×10^{-4}	1.56×10^{-4}	7.20×10^{-5}	2.46
-72.0	0.0	5.436×10^4	3.78×10^{-4}	3.78×10^{-4}	2.82×10^{-4}	1.30×10^{-4}	4.41
-67.0	-18.0	4.052×10^4	4.83×10^{-4}	4.81×10^{-4}	3.56×10^{-4}	1.63×10^{-4}	5.62
-60.0	-36.0	9.122×10^4	3.16×10^{-4}	3.15×10^{-4}	2.33×10^{-4}	1.07×10^{-4}	3.68
-60.0	36.0	3.583×10^4	4.93×10^{-4}	4.92×10^{-4}	3.62×10^{-4}	1.66×10^{-4}	5.74
-54.0	0.0	4.535×10^5	1.71×10^{-4}	1.70×10^{-4}	1.27×10^{-4}	5.86×10^{-5}	1.99
-45.0	18.0	2.543×10^5	1.85×10^{-4}	1.84×10^{-4}	1.35×10^{-4}	6.22×10^{-5}	2.14
-45.0	54.0	5.891×10^4	3.85×10^{-4}	3.85×10^{-4}	2.86×10^{-4}	1.32×10^{-4}	4.50
-36.0	0.0	5.314×10^5	1.32×10^{-4}	1.31×10^{-4}	9.64×10^{-5}	4.43×10^{-5}	1.53
-30.0	36.0	1.104×10^5	2.82×10^{-4}	2.81×10^{-4}	2.07×10^{-4}	9.51×10^{-5}	3.28
-23.0	18.0	3.586×10^5	1.53×10^{-4}	1.53×10^{-4}	1.12×10^{-4}	5.16×10^{-5}	1.79
-22.0	-18.0	1.914×10^5	2.12×10^{-4}	2.12×10^{-4}	1.55×10^{-4}	7.09×10^{-5}	2.47
-18.0	0.0	5.644×10^5	1.21×10^{-4}	1.20×10^{-4}	8.84×10^{-5}	4.06×10^{-5}	1.40
0.0	-18.0	2.853×10^5	1.74×10^{-4}	1.73×10^{-4}	1.27×10^{-4}	5.81×10^{-5}	2.02
0.0	0.0	1.165×10^6	9.57×10^{-5}	9.51×10^{-5}	6.97×10^{-5}	3.20×10^{-5}	1.11
0.0	18.0	4.576×10^5	1.39×10^{-4}	1.38×10^{-4}	1.01×10^{-4}	4.66×10^{-5}	1.61
0.0	36.0	1.371×10^5	2.59×10^{-4}	2.58×10^{-4}	1.91×10^{-4}	8.79×10^{-5}	3.02
0.0	72.0	2.891×10^4	5.42×10^{-4}	5.42×10^{-4}	4.01×10^{-4}	1.85×10^{-4}	6.33
18.0	0.0	4.222×10^5	1.53×10^{-4}	1.53×10^{-4}	1.12×10^{-4}	5.11×10^{-5}	1.79
23.0	-18.0	1.463×10^5	2.44×10^{-4}	2.45×10^{-4}	1.78×10^{-4}	8.12×10^{-5}	2.86
23.0	18.0	1.725×10^5	2.29×10^{-4}	2.30×10^{-4}	1.67×10^{-4}	7.62×10^{-5}	2.68
30.0	-36.0	2.759×10^4	9.86×10^{-4}	9.79×10^{-4}	7.25×10^{-4}	3.37×10^{-4}	11.43
36.0	0.0	8.999×10^5	1.00×10^{-4}	9.94×10^{-5}	7.33×10^{-5}	3.38×10^{-5}	1.16
37.0	90.0	3.360×10^5	1.58×10^{-4}	1.58×10^{-4}	1.17×10^{-4}	5.40×10^{-5}	1.85
45.0	-54.0	5.093×10^4	4.02×10^{-4}	4.03×10^{-4}	2.98×10^{-4}	1.38×10^{-4}	4.70
45.0	-18.0	6.642×10^4	3.56×10^{-4}	3.52×10^{-4}	2.61×10^{-4}	1.21×10^{-4}	4.11
54.0	0.0	3.763×10^5	1.51×10^{-4}	1.50×10^{-4}	1.11×10^{-4}	5.11×10^{-5}	1.75
68.0	18.0	1.551×10^5	2.43×10^{-4}	2.35×10^{-4}	1.76×10^{-4}	8.18×10^{-5}	2.74
72.0	0.0	2.961×10^5	2.11×10^{-4}	2.07×10^{-4}	1.54×10^{-4}	7.16×10^{-5}	2.41
90.0	-54.0	1.589×10^4	1.23×10^{-3}	1.23×10^{-3}	9.24×10^{-4}	4.28×10^{-4}	14.32
90.0	-36.0	1.589×10^4	1.23×10^{-3}	1.23×10^{-3}	9.24×10^{-4}	4.28×10^{-4}	14.32
90.0	-18.0	1.979×10^4	6.50×10^{-4}	6.39×10^{-4}	4.77×10^{-4}	2.21×10^{-4}	7.46
90.0	0.0	2.472×10^5	1.86×10^{-4}	1.82×10^{-4}	1.35×10^{-4}	6.30×10^{-5}	2.12

Table 5—Continued

l	b	Exposure (sec)	100 keV ^a	511 keV ^a	1000 keV ^a	5000 keV ^a	Scale Factor ^b
90.0	54.0	4.553×10^4	4.21×10^{-4}	4.21×10^{-4}	3.11×10^{-4}	1.43×10^{-4}	4.92
90.0	72.0	1.462×10^5	2.56×10^{-4}	2.57×10^{-4}	1.89×10^{-4}	8.71×10^{-5}	2.99
108.0	0.0	1.783×10^5	2.16×10^{-4}	2.16×10^{-4}	1.59×10^{-4}	7.35×10^{-5}	2.52
126.0	0.0	1.257×10^5	2.61×10^{-4}	2.60×10^{-4}	1.92×10^{-4}	8.84×10^{-5}	3.04
135.0	54.0	6.624×10^4	3.50×10^{-4}	3.52×10^{-4}	2.59×10^{-4}	1.19×10^{-4}	4.10
144.0	0.0	1.746×10^4	7.06×10^{-4}	7.00×10^{-4}	5.16×10^{-4}	2.38×10^{-4}	8.17
162.0	0.0	7.835×10^4	4.97×10^{-4}	4.92×10^{-4}	3.63×10^{-4}	1.67×10^{-4}	5.74

^aNarrow line sensitivity (photons cm⁻² sec⁻¹)

^bScale factor to apply to sensitivity plot of Fig. 7.

Table 6. Point Source Sensitivities - LMXB

Source	Exposure (sec)	50 keV ^a	100 keV ^a	511 keV ^a	1000 keV ^a	Scale Factor ^b
1A 1742-294	4.270×10^6	1.32×10^{-3}	1.66×10^{-3}	8.36×10^{-4}	6.55×10^{-4}	5.33
1E 1740	4.282×10^6	9.56×10^{-4}	1.18×10^{-3}	6.00×10^{-4}	4.62×10^{-4}	3.76
3A 1728-169	4.426×10^6	1.69×10^{-4}	2.11×10^{-4}	1.09×10^{-4}	8.44×10^{-5}	0.69
3A 1822-371	4.346×10^6	1.91×10^{-4}	2.29×10^{-4}	1.19×10^{-4}	9.33×10^{-5}	0.76
4U 1626-67	8.267×10^5	4.66×10^{-4}	5.47×10^{-4}	2.90×10^{-4}	2.54×10^{-4}	2.07
4U 1730-335	4.181×10^6	3.17×10^{-4}	3.47×10^{-4}	2.07×10^{-4}	1.61×10^{-4}	1.31
4U 1735-444	4.236×10^6	2.05×10^{-4}	2.44×10^{-4}	1.21×10^{-4}	1.02×10^{-4}	0.83
4U 1916-053	1.552×10^6	2.53×10^{-4}	2.98×10^{-4}	1.61×10^{-4}	1.47×10^{-4}	1.19
Aql X-1	1.073×10^6	2.58×10^{-4}	3.47×10^{-4}	1.78×10^{-4}	1.64×10^{-4}	1.34
Cir X-1	1.853×10^6	2.60×10^{-4}	3.44×10^{-4}	1.71×10^{-4}	1.39×10^{-4}	1.14
Cyg X-2	5.252×10^5	4.32×10^{-4}	5.18×10^{-4}	2.50×10^{-4}	2.06×10^{-4}	1.68
EXO 0748-676	7.796×10^5	2.29×10^{-4}	2.92×10^{-4}	1.41×10^{-4}	1.27×10^{-4}	1.04
GRO J1655-40	4.052×10^6	1.84×10^{-4}	2.37×10^{-4}	1.21×10^{-4}	1.01×10^{-4}	0.82
GRS 1739-278	4.196×10^6	2.11×10^{-4}	2.57×10^{-4}	1.33×10^{-4}	1.03×10^{-4}	0.84
GX 1+4	4.196×10^6	1.41×10^{-4}	1.82×10^{-4}	9.49×10^{-5}	7.36×10^{-5}	0.60
GX 13+1	4.060×10^6	1.79×10^{-4}	2.40×10^{-4}	1.25×10^{-4}	9.93×10^{-5}	0.81
GX 17+2	3.732×10^6	2.12×10^{-4}	2.59×10^{-4}	1.35×10^{-4}	1.08×10^{-4}	0.88
GX 3+1	4.196×10^6	1.59×10^{-4}	1.94×10^{-4}	1.01×10^{-4}	7.86×10^{-5}	0.64
GX 349+2	4.320×10^6	1.79×10^{-4}	2.28×10^{-4}	1.16×10^{-4}	9.28×10^{-5}	0.76
GX 354-0	4.354×10^6	3.15×10^{-4}	3.95×10^{-4}	2.09×10^{-4}	1.62×10^{-4}	1.32
GX 5-1	4.318×10^6	1.43×10^{-4}	1.90×10^{-4}	9.83×10^{-5}	7.62×10^{-5}	0.62
GX 9+1	4.116×10^6	2.31×10^{-4}	2.87×10^{-4}	1.49×10^{-4}	1.16×10^{-4}	0.94
Ginga 1826-24	4.116×10^6	1.65×10^{-4}	2.10×10^{-4}	1.09×10^{-4}	8.47×10^{-5}	0.69
H 0614+091	6.176×10^5	4.86×10^{-4}	5.35×10^{-4}	2.54×10^{-4}	1.84×10^{-4}	1.50
H 1608-522	1.708×10^6	2.23×10^{-4}	2.95×10^{-4}	1.47×10^{-4}	1.36×10^{-4}	1.11
H 1636-536	1.559×10^6	2.32×10^{-4}	2.85×10^{-4}	1.54×10^{-4}	1.39×10^{-4}	1.13
H 1702-429	4.062×10^6	2.55×10^{-4}	3.10×10^{-4}	1.63×10^{-4}	1.40×10^{-4}	1.14
H 1705-250	4.761×10^6	1.55×10^{-4}	1.90×10^{-4}	9.84×10^{-5}	7.62×10^{-5}	0.62
H 1705-440	3.859×10^6	2.43×10^{-4}	3.09×10^{-4}	1.56×10^{-4}	1.37×10^{-4}	1.11
H 1820-303	4.141×10^6	1.55×10^{-4}	1.96×10^{-4}	1.00×10^{-4}	7.66×10^{-5}	0.62
IGR J16358-4726	2.066×10^6	3.87×10^{-4}	4.79×10^{-4}	2.46×10^{-4}	2.24×10^{-4}	1.83
IGR J16418-4532	1.915×10^6	2.73×10^{-4}	3.50×10^{-4}	1.77×10^{-4}	1.61×10^{-4}	1.31
IGR J17464-3213	4.455×10^6	1.47×10^{-4}	1.96×10^{-4}	1.02×10^{-4}	8.07×10^{-5}	0.66
KS 1741-293	4.270×10^6	9.17×10^{-3}	1.14×10^{-2}	5.70×10^{-3}	4.52×10^{-3}	36.83
Sco X-1	6.764×10^5	2.78×10^{-4}	3.79×10^{-4}	2.00×10^{-4}	1.47×10^{-4}	1.20
Ser X-1	1.413×10^6	2.32×10^{-4}	3.01×10^{-4}	1.52×10^{-4}	1.39×10^{-4}	1.14
XTE J1748-288	4.196×10^6	2.78×10^{-4}	3.33×10^{-4}	1.73×10^{-4}	1.32×10^{-4}	1.08
XTE J1807-294	4.187×10^6	1.44×10^{-4}	1.87×10^{-4}	9.65×10^{-5}	7.43×10^{-5}	0.61
LMXB Average	1.219×10^8	3.47×10^{-5}	4.36×10^{-5}	2.26×10^{-5}	1.82×10^{-5}	0.15

^aNarrow line sensitivity (photons cm⁻² sec⁻¹)

^bScale factor to apply to sensitivity plot of Fig. 10.

Table 7. Point Source Sensitivities - HMXB

Source	Exposure (sec)	50 keV ^a	100 keV ^a	511 keV ^a	1000 keV ^a	Scale Factor ^b
1E 1145	8.413×10^5	3.28×10^{-4}	4.42×10^{-4}	2.21×10^{-4}	1.75×10^{-4}	1.43
3A 0114+650	3.109×10^5	8.06×10^{-4}	1.05×10^{-3}	5.26×10^{-4}	4.43×10^{-4}	3.61
3A 2206+543	4.375×10^5	5.14×10^{-4}	6.53×10^{-4}	3.32×10^{-4}	2.85×10^{-4}	2.32
4U 1036-56	2.016×10^5	7.82×10^{-4}	9.98×10^{-4}	4.97×10^{-4}	4.35×10^{-4}	3.54
4U 1700-377	4.332×10^6	1.78×10^{-4}	2.33×10^{-4}	1.20×10^{-4}	9.72×10^{-5}	0.79
4U 1901+03	1.494×10^6	2.22×10^{-4}	2.91×10^{-4}	1.48×10^{-4}	1.35×10^{-4}	1.10
Cen X-3	3.955×10^5	3.81×10^{-4}	5.06×10^{-4}	2.59×10^{-4}	2.30×10^{-4}	1.87
Cyg X-1	1.417×10^6	1.75×10^{-4}	2.21×10^{-4}	1.11×10^{-4}	9.34×10^{-5}	0.76
GX 301-2	8.190×10^5	2.93×10^{-4}	3.84×10^{-4}	1.93×10^{-4}	1.46×10^{-4}	1.19
Ginga 1843+009	1.494×10^6	2.16×10^{-4}	2.77×10^{-4}	1.43×10^{-4}	1.30×10^{-4}	1.06
H 0115+634	3.109×10^5	7.98×10^{-4}	1.03×10^{-3}	5.14×10^{-4}	4.33×10^{-4}	3.53
H 1538-522	1.723×10^6	2.39×10^{-4}	3.08×10^{-4}	1.54×10^{-4}	1.37×10^{-4}	1.12
H 1907+097	1.123×10^6	6.69×10^{-4}	8.83×10^{-4}	4.47×10^{-4}	4.14×10^{-4}	3.37
IGR J16318-4848	1.759×10^6	3.58×10^{-4}	4.72×10^{-4}	2.36×10^{-4}	2.16×10^{-4}	1.76
IGR J16320-4751	1.980×10^6	4.81×10^{-4}	6.21×10^{-4}	3.10×10^{-4}	2.82×10^{-4}	2.29
IGR J17391-3021	4.463×10^6	2.28×10^{-4}	2.83×10^{-4}	1.46×10^{-4}	1.13×10^{-4}	0.92
KS 1947+300	1.188×10^6	2.16×10^{-4}	2.74×10^{-4}	1.38×10^{-4}	1.14×10^{-4}	0.93
LMC X-4	7.796×10^5	2.27×10^{-4}	3.07×10^{-4}	1.56×10^{-4}	1.40×10^{-4}	1.14
OA0 1657-415	4.052×10^6	2.05×10^{-4}	2.64×10^{-4}	1.32×10^{-4}	1.15×10^{-4}	0.94
RX J0053	1.077×10^5	1.29×10^{-3}	1.53×10^{-3}	7.80×10^{-4}	5.51×10^{-4}	4.49
SAX J2103	1.448×10^6	2.48×10^{-4}	2.96×10^{-4}	1.42×10^{-4}	1.19×10^{-4}	0.97
SMC X-1	1.077×10^5	1.04×10^{-3}	1.36×10^{-3}	6.94×10^{-4}	5.01×10^{-4}	4.09
Vela X-1	1.043×10^6	1.99×10^{-4}	2.63×10^{-4}	1.37×10^{-4}	1.08×10^{-4}	0.88
XTE J1855-026	1.593×10^6	2.39×10^{-4}	3.09×10^{-4}	1.54×10^{-4}	1.42×10^{-4}	1.16
XTE J1908+094	4.392×10^5	1.03×10^{-3}	1.32×10^{-3}	6.71×10^{-4}	6.25×10^{-4}	5.10
HMXB Average	3.386×10^7	5.59×10^{-5}	7.21×10^{-5}	3.64×10^{-5}	3.08×10^{-5}	0.25

^aNarrow line sensitivity [(photons cm⁻² sec⁻¹)

^bScale factor to apply to sensitivity plot of Fig. 10.

Table 8. Point Source Sensitivities - Pulsars

Source	Exposure (sec)	50 keV ^a	100 keV ^a	511 keV ^a	1000 keV ^a	Scale Factor ^b
PSR B1509-58	1.591×10^6	2.63×10^{-4}	3.43×10^{-4}	1.75×10^{-4}	1.39×10^{-4}	1.13
Vela Pulsar	1.043×10^6	1.84×10^{-4}	2.46×10^{-4}	1.27×10^{-4}	1.01×10^{-4}	0.82
Pulsar Average	2.634×10^6	1.51×10^{-4}	2.00×10^{-4}	1.03×10^{-4}	8.14×10^{-5}	0.66

^aNarrow line sensitivity (photons cm⁻² sec⁻¹)

^bScale factor to apply to sensitivity plot of Fig. 10.

Table 9. Point Source Sensitivities - Quasars

Source	Exposure (sec)	50 keV ^a	100 keV ^a	511 keV ^a	1000 keV ^a	Scale Factor ^b
3C 279	3.672×10^5	3.31×10^{-4}	4.38×10^{-4}	2.21×10^{-4}	1.67×10^{-4}	1.36
PKS 0637-752	8.873×10^5	2.31×10^{-4}	3.12×10^{-4}	1.58×10^{-4}	1.40×10^{-4}	1.14
PKS 1622-297	3.838×10^6	2.51×10^{-4}	2.96×10^{-4}	1.47×10^{-4}	1.18×10^{-4}	0.96
PKS 1830-211	4.165×10^6	1.83×10^{-4}	2.25×10^{-4}	1.19×10^{-4}	9.32×10^{-5}	0.76

^aNarrow line sensitivity (photons cm⁻² sec⁻¹)

^bScale factor to apply to sensitivity plot of Fig. 10.

Table 10. Point Source Sensitivities - Supernova Remnants

Source	Exposure (sec)	50 keV ^a	100 keV ^a	511 keV ^a	1000 keV ^a	Scale Factor ^b
Cas A	3.659×10^5	5.02×10^{-4}	6.63×10^{-4}	3.32×10^{-4}	2.85×10^{-4}	2.32
Crab	7.108×10^5	2.47×10^{-4}	3.09×10^{-4}	1.57×10^{-4}	1.23×10^{-4}	1.00
SNR Average	1.077×10^6	2.22×10^{-4}	2.80×10^{-4}	1.42×10^{-4}	1.13×10^{-4}	0.92

^aNarrow line sensitivity (photons cm⁻² sec⁻¹)

^bScale factor to apply to sensitivity plot of Fig. 10.

Table 11. Point Source Sensitivities - Seyferts

Source	Exposure (sec)	50 keV ^a	100 keV ^a	511 keV ^a	1000 keV ^a	Scale Factor ^b
1H 1934-063	1.070×10^6	3.16×10^{-4}	3.88×10^{-4}	1.92×10^{-4}	1.70×10^{-4}	1.38
3C 111	8.860×10^4	1.91×10^{-3}	2.28×10^{-3}	1.11×10^{-3}	1.02×10^{-3}	8.34
Cen A	3.050×10^5	4.08×10^{-4}	5.14×10^{-4}	2.61×10^{-4}	2.13×10^{-4}	1.73
Circinus Galaxy	8.460×10^5	2.73×10^{-4}	3.66×10^{-4}	1.85×10^{-4}	1.39×10^{-4}	1.14
IC 4329A	3.050×10^5	3.48×10^{-4}	4.67×10^{-4}	2.40×10^{-4}	1.87×10^{-4}	1.53
IC 4518	1.589×10^6	3.09×10^{-4}	3.82×10^{-4}	1.89×10^{-4}	1.63×10^{-4}	1.33
IGR J18027-1455	3.938×10^6	1.97×10^{-4}	2.45×10^{-4}	1.24×10^{-4}	9.66×10^{-5}	0.79
MCG+08-11-011	9.731×10^4	1.31×10^{-3}	1.72×10^{-3}	8.67×10^{-4}	7.49×10^{-4}	6.10
MCG-05-23-016	3.452×10^5	8.75×10^{-4}	8.19×10^{-4}	3.78×10^{-4}	2.49×10^{-4}	2.03
MCG-06-30-015	3.050×10^5	3.46×10^{-4}	4.57×10^{-4}	2.34×10^{-4}	1.84×10^{-4}	1.50
MR2251-175	1.335×10^5	4.65×10^{-4}	6.33×10^{-4}	3.22×10^{-4}	2.83×10^{-4}	2.30
NGC 1275	4.132×10^4	3.18×10^{-3}	3.92×10^{-3}	1.93×10^{-3}	1.75×10^{-3}	14.27
NGC 4388	6.114×10^5	3.73×10^{-4}	4.47×10^{-4}	2.18×10^{-4}	1.67×10^{-4}	1.36
NGC 4507	9.870×10^4	6.85×10^{-4}	9.52×10^{-4}	4.90×10^{-4}	4.75×10^{-4}	3.87
NGC 4736	6.334×10^5	3.23×10^{-4}	4.11×10^{-4}	2.02×10^{-4}	1.78×10^{-4}	1.45
NGC 4945	5.932×10^5	3.87×10^{-4}	4.93×10^{-4}	2.46×10^{-4}	1.85×10^{-4}	1.51
NGC 5033	4.935×10^5	2.98×10^{-4}	4.13×10^{-4}	2.10×10^{-4}	1.86×10^{-4}	1.52
NGC 6300	1.280×10^6	3.84×10^{-4}	4.55×10^{-4}	2.33×10^{-4}	2.10×10^{-4}	1.71
NGC 6814	8.316×10^5	4.41×10^{-4}	5.35×10^{-4}	2.64×10^{-4}	2.30×10^{-4}	1.88
Seyfert Average	1.361×10^7	7.23×10^{-5}	7.25×10^{-5}	4.23×10^{-5}	4.38×10^{-5}	0.36

^aNarrow line sensitivity (photons $\text{cm}^{-2} \text{sec}^{-1}$)

^bScale factor to apply to sensitivity plot of Fig. 10.

Table 12. Point Source Sensitivities - Miscellaneous/Unidentified

Source	Exposure (sec)	50 keV ^a	100 keV ^a	511 keV ^a	1000 keV ^a	Scale Factor ^b
2RXP J130159	7.424×10^5	2.77×10^{-4}	3.74×10^{-4}	1.91×10^{-4}	1.41×10^{-4}	1.15
3A 1850-087	1.762×10^6	2.55×10^{-4}	3.32×10^{-4}	1.67×10^{-4}	1.47×10^{-4}	1.20
3C 273	3.672×10^5	2.82×10^{-4}	3.34×10^{-4}	1.95×10^{-4}	1.50×10^{-4}	1.22
Coma Cluster	4.935×10^5	2.74×10^{-4}	3.68×10^{-4}	1.87×10^{-4}	1.66×10^{-4}	1.35
Cyg X-3	1.448×10^6	1.84×10^{-4}	2.25×10^{-4}	1.20×10^{-4}	9.99×10^{-5}	0.81
IGR J00291+5934	2.894×10^5	2.44×10^{-3}	3.08×10^{-3}	1.58×10^{-3}	1.35×10^{-3}	11.04
IGR J00370+6122	2.894×10^5	6.05×10^{-4}	8.06×10^{-4}	4.07×10^{-4}	3.49×10^{-4}	2.85
IGR J06074+2205	7.108×10^5	2.67×10^{-4}	3.42×10^{-4}	1.74×10^{-4}	1.37×10^{-4}	1.12
IGR J07597-3842	1.071×10^6	2.38×10^{-4}	2.86×10^{-4}	1.52×10^{-4}	1.20×10^{-4}	0.98
IGR J11305-6256	8.413×10^5	3.65×10^{-4}	4.66×10^{-4}	2.28×10^{-4}	1.97×10^{-4}	1.61
IGR J16167-4957	1.799×10^6	5.47×10^{-4}	6.61×10^{-4}	3.51×10^{-4}	3.23×10^{-4}	2.63
IGR J16195-4945	1.799×10^6	5.43×10^{-4}	6.98×10^{-4}	3.47×10^{-4}	3.19×10^{-4}	2.60
IGR J16479-4514	2.266×10^6	2.57×10^{-4}	3.30×10^{-4}	1.66×10^{-4}	1.52×10^{-4}	1.24
IGR J16558-5203	1.641×10^6	2.22×10^{-4}	2.85×10^{-4}	1.46×10^{-4}	1.32×10^{-4}	1.08
IGR J17252-3616	4.103×10^6	1.57×10^{-4}	2.06×10^{-4}	1.07×10^{-4}	8.51×10^{-5}	0.69
IGR J17303-0601	1.189×10^6	3.51×10^{-4}	4.49×10^{-4}	2.25×10^{-4}	1.92×10^{-4}	1.57
IGR J17597-2201	4.278×10^6	1.61×10^{-4}	2.13×10^{-4}	1.12×10^{-4}	8.66×10^{-5}	0.71
IGR J18325-0756	1.737×10^6	2.42×10^{-4}	3.10×10^{-4}	1.58×10^{-4}	1.36×10^{-4}	1.11
IGR J18450-0435	1.552×10^6	2.60×10^{-4}	3.42×10^{-4}	1.74×10^{-4}	1.55×10^{-4}	1.27
IGR J18483-0311	1.873×10^6	2.61×10^{-4}	3.38×10^{-4}	1.71×10^{-4}	1.54×10^{-4}	1.26
IGR J19140+0951	1.123×10^6	3.00×10^{-4}	3.72×10^{-4}	1.99×10^{-4}	1.84×10^{-4}	1.50
IGR J21247+5058	6.295×10^5	3.62×10^{-4}	4.63×10^{-4}	2.31×10^{-4}	1.96×10^{-4}	1.60
RX J0146	3.109×10^5	7.73×10^{-4}	9.34×10^{-4}	4.67×10^{-4}	3.91×10^{-4}	3.18
SAX J1744	4.270×10^6	7.38×10^{-3}	9.29×10^{-3}	4.70×10^{-3}	3.69×10^{-3}	30.09
SAX J1805	4.116×10^6	2.32×10^{-4}	2.90×10^{-4}	1.47×10^{-4}	1.14×10^{-4}	0.93
SS 433	1.261×10^6	2.38×10^{-4}	3.13×10^{-4}	1.59×10^{-4}	1.47×10^{-4}	1.20
V709 Cas	2.894×10^5	2.52×10^{-3}	3.29×10^{-3}	1.63×10^{-3}	1.39×10^{-3}	11.32
XTE J1901+014	7.882×10^5	3.20×10^{-4}	4.20×10^{-4}	2.14×10^{-4}	1.94×10^{-4}	1.58
γ Cas	2.636×10^5	6.63×10^{-4}	8.63×10^{-4}	4.33×10^{-4}	3.60×10^{-4}	2.94

^aNarrow line sensitivity (photons $\text{cm}^{-2} \text{sec}^{-1}$)

^bScale factor to apply to sensitivity plot of Fig. 10.