Rejuvenated radio galaxies J0041+3224 and J1835+6204: how long can the quiescent phase of nuclear activity last?

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ABSTRACT

We present radio observations of two well-known double-double radio galaxies, J0041+3224 and J1835+6204, at frequencies ranging from 150 to 8460 MHz, using both the Giant Metrewave Radio Telescope and the Very Large Array. These observations, over a large radio frequency range, enable us to determine the spectra of the inner and outer lobes. Our detailed spectral ageing analysis of their inner and outer lobes demonstrates that the outer doubles of double-double radio galaxies are created by the previous cycle of activity, while the inner doubles are due to the present cycle of activity. The (core subtracted) spectra of the inner doubles of both sources are power laws over a large frequency range. We found that the duration of the quiescent phase of J0041+3224 is between 4 and 28 per cent of the active phase of the previous activity. The outer north-western lobe of J1835+6204 has a compact hotspot and the regions of both the outer hotspots have close to power-law (rather than curved) spectra, which indicates that the outer lobes are still fed by jet material ejected in the previous episode just before the central engine stopped powering the jet. We estimate that the duration of the quiescent phase of J1835+6204 is $\lesssim$5 per cent of the duration of the active phase of the previous activity. Therefore, we conclude that the duration of the quiescent phase can be as short as a few per cent of the active phase in radio galaxies of this type.

Key words: galaxies: active – galaxies: nuclei – galaxies: individual: J0041+3224, J1835+6204 – radio continuum: galaxies

1 INTRODUCTION

It is established beyond doubt that jet forming activity in radio galaxies is episodic in nature. This episodic jet-forming activity often gives rise to Double-Double Radio Galaxies (DDRGs), which are defined to be those having a pair of double radio sources with a common centre, and are thought to occur when a new epoch of jet activity is triggered in a radio galaxy with older lobes still visible from the activity of the previous epoch (Schoenmakers et al. 2000a). At present, about two dozen DDRGs are known in the literature (see Saikia & Jamrozy 2009 for a review). Although the DDRGs are more common, in fact more than two episodes of jet forming activity are physically plausible: Brocksopp et al. (2007) reported the first example of three such episodes in a

radio galaxy, and recently Hota et al. (2011) reported a possible case of a triple-double radio galaxy dubbed 'SPECA' (J1409−0302).

The very existence of well-shaped and well confined radio galaxy lobes suggests that there must be some medium confining the lobes either through ram pressure or thermal pressure. Even radio galaxies whose optical environments are known to be comparatively poor are found to possess X-ray-emitting hot gas environments, corresponding to group or cluster-scale X-ray luminosities, whose gas properties can now be modelled in detail (see Hardcastle & Worrall, 1999; Hardcastle & Worrall, 2000; Worrall & Birkinshaw, 2000; Croston et al., 2003; Croston et al., 2004; Evans et al., 2005; Belsole et al., 2007; Croston et al., 2008; Konar et al., 2009) This environment emits via thermal bremsstrahlung in the X-ray band, and can be classified as a poor-cluster to group scale environment. The magnetised relativistic plasma

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(MRP) of the radio lobes does not mix well with the external thermal gas in the environment, as can be inferred from limits on internal depolarization of lobes at low frequencies, nor does it diffuse out completely into the IntraCluster Medium (ICM) even long after the jet stops feeding the lobes. Kaiser, Schoenmakers & Röttgering (2000) argued against the entrainment of material from the surrounding hot-gas environment through the lobe periphery. However, the observed well-confined inner lobes of DDRGs are not capable of existing in such a tenuous cocoon of matter deposited in the previous cycle of jet activity (the typical number density of relativistic radiating particles is $\sim 10^{-10}$ cm^{-3}). Therefore, Kaiser et al. (2000) proposed that the ambient thermal medium is a two phase medium. One phase consists of dense warm clouds and the other phase is the hot-gas medium, which has the characteristics of the general ICM, with a temperature of $\sim 10^7$K and a central particle number density of $10^{-1} - 10^{-3}$ cm^{-3}. The dense, warm cloud phase has a temperature of $\sim 10^4$K and a particle number density of ~ 100 cm^{-3}, with a very low filling factor of $10^{-8} - 10^{-3}$, dependent on distance from the Active Galactic Nucleus (AGN). The hot-gas medium has a filling factor ~ 1. Kaiser et al. (2000) discussed in detail, in Section 2 of their paper, the observational evidence for these warm clouds through optical emission line detections. The warm clouds are dynamically unimportant, as their filling factor is extremely small. When the radio lobes expand very rapidly, these clouds are not efficiently accelerated by the lobe expansion, being heavier than the hot-gas phase, and so are overtaken by the contact discontinuity. This means that the warm clouds become dispersed inside the lobe and provide the required matter density to prevent the inner jets from being ballistic. In this model, to understand the dynamics of these sources it is necessary to know the pressure and matter density internal and external to the lobes.

The condition for a radio galaxy to restart is not well understood. Schoenmakers et al. (2000a) explored the possibility of causing restarting jets in DDRGs by interaction and merger events, although there is no observational evidence for this model. Numerical simulations of colliding galaxies show that these usually do not merge completely in the first encounter (Barnes & Hernquist, 1996). According to this scenario, the in-falling galaxy loses a large fraction of its gas and stars to the main galaxy, while parts of it pass through the main galaxy. These parts will turn around and collide again with the main galaxy. If the first encounter triggers the AGN, then the subsequent encounters destabilise it and jet interruption occurs. However, a problem with this scenario is that, as argued by Schoenmakers et al. (2000a), it would lead to the expectation of finding many more restarting radio galaxies than we see today. The mechanism for triggering multiple episodes of jet activity could also be closely related to the feedback mechanism proposed for solving the cooling flow problem in clusters of galaxies (see McNamara & Nulsen, 2007). Recently, the study of Jetha et al. (2008) showed that the restarting phenomenon in B2 0838+32A is consistent with feedback-driven accretion by the AGN. Therefore, studying DDRGs is very important to understand the conditions for the phenomenon of restarting jets, their relation to feedback mechanisms, and the dynamics of radio galaxies in general.

To understand the dynamics of radio galaxies we need to know the injection spectral index (which is related to the particle acceleration process), the spectral age (which is a good approximation to dynamical age), the jet power (which is connected to the accretion rate at the central engine), the magnetic field strength based on the equipartition assumption (which we know is reasonably accurate in FR II radio galaxies: Croston et al. 2005), and the internal pressure of the lobes (which helps us determine the composition of the relativistic particles, by comparing with the external pressure obtained from X-ray observations); all these parameters can be estimated from radio observations. Therefore, comprehensive radio studies of such sources are essential to pave the way towards a complete understanding of the dynamics of radio galaxies. **Our further work related to the time-scale of quiescent phase of episodic radio galaxies, injection index in different episodes and particle acceleration at hotspots, and the dynamics of DDRGs through radio and X-ray observations will be published soon by Konar et al. (in prep); Konar & Hardcastle (in prep) and Konar et al. (in prep) respectively.**

In this paper, we have carried out a multifrequency radio study of two DDRGs, namely J0041+3224 and J1835+6204. We present the Giant Meterwave Radio Telescope (GMRT) and Very Large Array (VLA) observations in Section 2, our observational results in Section 3, our spectral ageing analysis in Section 4, a discussion in Section 5 and concluding remarks in Section 6.

The cosmological parameters that we have used are $H_o = 71$ km s^{-1} Mpc^{-1}, $\Omega_M = 0.27$ and $\Omega_{vac} = 0.73$ (Spergel et al. 2003). In this cosmology, 1 arcsec corresponds to 5.733 kpc for the source J0041+3224 situated at an (estimated, see below) luminosity distance $D_L = 2486.3$ Mpc. For J1835+6204, 1 arcsec corresponds to 6.207 kpc based on its luminosity distance of $D_L = 2955.7$ Mpc.

2   OBSERVATIONS AND DATA REDUCTION

The images published in this paper are from the GMRT and the VLA. All GMRT data are from our observations with project code 10CKa01. **All VLA data are public data from the VLA archive.** The details of the observations with both telescopes are given in Table 1.

The shortest baseline for the C and D configuations of the VLA is 35 m. Our target sources, J0041+3224 and J1835+6204 have sizes ~ 2.8 arcmin and ~ 3.7 arcmin respectively; however, in both sources, at high frequencies, no single structure with dominant flux density is larger than 2 arcmin. As all the VLA maps (published by Saikia et al. 2006) used for flux density measurements of J0041+3224 were made from C-array data, where the largest angular size that can be mapped without loss of flux is 3 arcmin, our measurements for J0041+3224 should not be susceptible to loss of flux due to lack of short spacings in the uv-coverage. Alhough the 8.4 GHz map of J1835+6204 (Figure 2) looks continuous from hotspot to hotspot, this is (at least partially) due to the effects of comparatively poor resolution, which is why we do not see such continuity in the 4.8 GHz images at somewhat higher resolution. The 8.4-GHz image of J1835+6204 is made by us from a combination of observations made separately with the B and D configurations,

which again can image 3 arcmin regions without loss of flux. So this source is also not susceptible to any major loss of flux due to the lack of short spacings in the uv coverage. The shortest baseline of the GMRT is 100 m, and so we can safely map emission regions of a size of 7 arcmin at 1280 MHz without loss of flux; therefore, the GMRT can map even larger structure at any frequencies below 1280 MHz without losing any flux. We conclude that neither the VLA nor the GMRT images suffer from any significant loss of flux due to the lack of short spacings. In addition, we see no observational signatures of such flux losses, and so conclude that our flux measurements and spectral analysis is reliable.

2.1 GMRT observations

The observations were made in the standard manner, with each observation of the target source interspersed with observations of calibrators. For an observing run of any target source, one of 3C48, 3C147 or 3C286 was observed as a flux density and bandpass calibrator at a given frequency. At each frequency the source was observed in a full-synthesis run of approximately 9 hours including calibration overheads. On-source durations vary from 240−371 min. Full details of the GMRT array can be found at `http://www.gmrt.ncra.tifr.res.in`. The data collected were calibrated and reduced in the standard way using the NRAO `AIPS` software package. Flux density calibration of our GMRT data uses the scale of Baars et al. (1977). All of the GMRT data below L band were affected by Radio Frequency Interference (RFI) to various extents, and were extensively flagged by hand within `AIPS` to reduce the effects of RFI as much as possible.

2.2 VLA observations

The VLA data for the sources at various frequencies (∼1.4, 4.8, 8.4 GHz) were analysed to produce images. The data from our own observations and those from the archive are indicated in the observation log table (Table 1). All VLA observations were in snap-shot mode. The on-source observing time was between 7 to 251 min (see Table 1). One of the calibrators 3C48, 3C147 or 3C286 was observed as a flux density calibrator at each frequency of a given observing run. As in the case of the GMRT data, the VLA data were edited and reduced using the `AIPS` package. The use of VLA archival data to supplement our GMRT observations was essential to constrain the spectra of the DDRGs and to allow us to carry out the spectral ageing analysis. All flux densities are on the Baars et al. (1977) scale.

3 OBSERVATIONAL RESULTS

3.1 Observed parameters and structures

The images of the entire source using the GMRT and the VLA are presented in Figures 1 and 2, while the observational parameters and some of the observed properties are presented in Table 2 and 3. The flux densities of different components were estimated by specifying an area around each component. **All the flux density measurements have been done with the AIPS task 'TVSTAT' which**
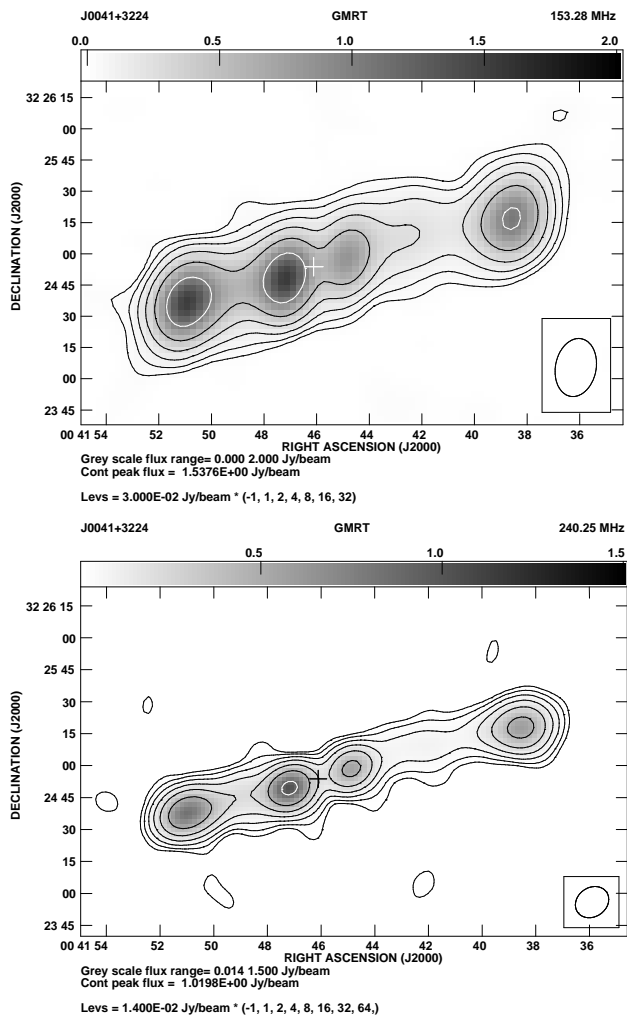

Figure 1. Full resolution radio images of J0041+3224 are shown. The frequency of each image and the telescope with which the image is made are given at the top of each image. The peak flux density, greyscale level, first contour and the contour levels are all given at the bottom of each image. A + sign indicates the position of the optical galaxy. We have not shown the 325 MHz image as its quality is not good enough. However, the flux density seems to be consistent when compared with all other frequencies. The rms noises and beam sizes are given in Table 2.

requires selection of a shaded polygonal area in the region of interest. 'TVSTAT' adds up all the pixel values within the shaded area to give the total flux density of the shaded area. We measured the flux density of each component using several different areas around the radio emitting component, which gives an idea of the systematic errors due to our choice of region. The change in flux density of each component, when measured this way, is only a few per cent. Flux density values that have been measured directly by us from the FITS maps with **the AIPS task 'TVSTAT'** are assumed to have the following error values: 5 per cent for 1400, 4860 and 8460 MHz VLA measurements; 7 per cent for 1287 and 610 MHz GMRT measurements; and 15 per cent for 332, 240 and 153 MHz GMRT measurements. Other flux density values quoted in the pa-
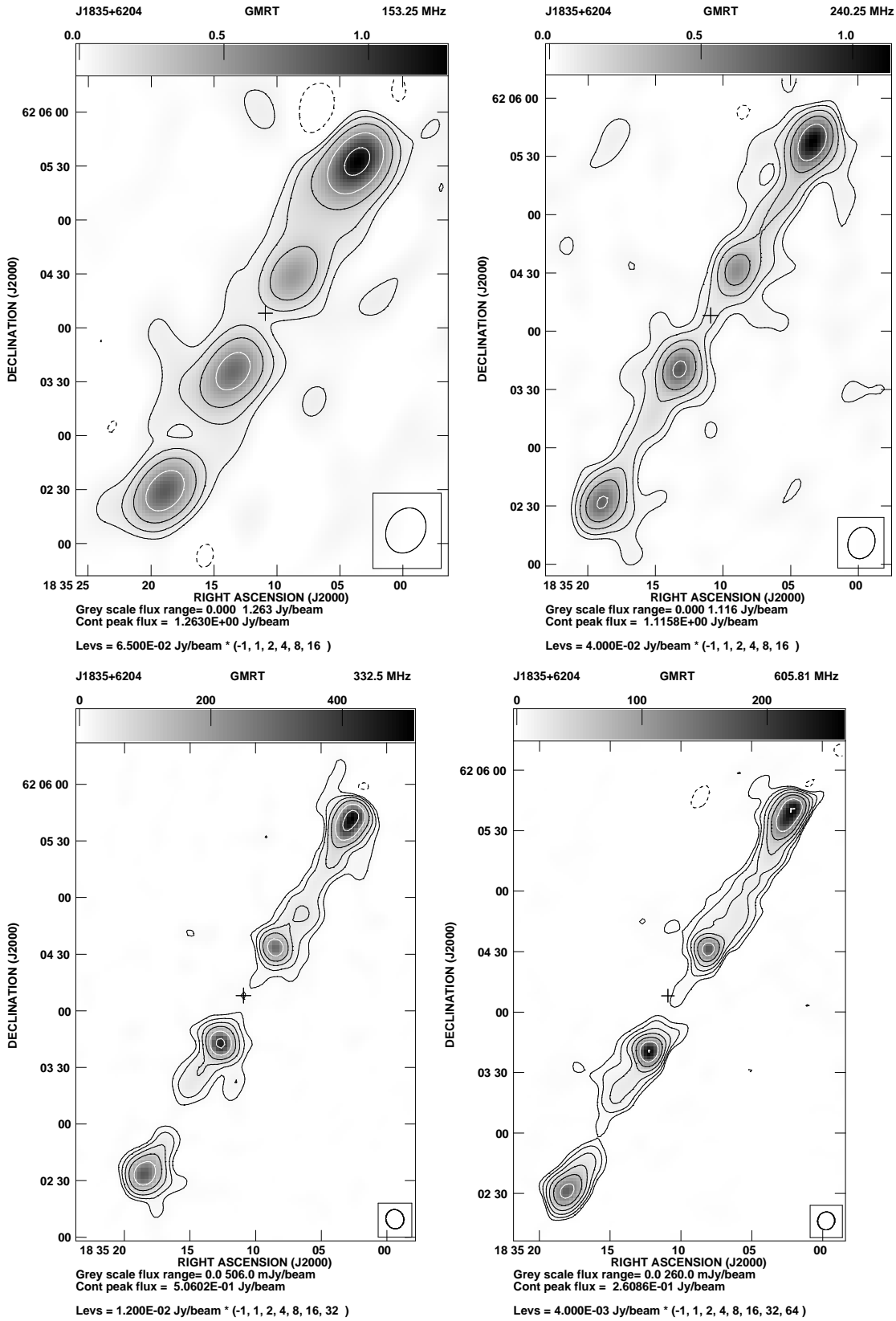
Figure 2. Full resolution radio images of J1835+6204 are shown. The exact frequency, and the name of the telescope (and observation codes of the VLA data sets used to produce the images) are given at the top of each image. Some of the VLA images are produced by combining two data sets to detect the diffuse emission. The peak flux density, greyscale level, 1st contour and the contour levels are all given at the bottom of each image. A + sign indicates the position of the optical galaxy. The rms noises and the beam sizes are given in Table 3.
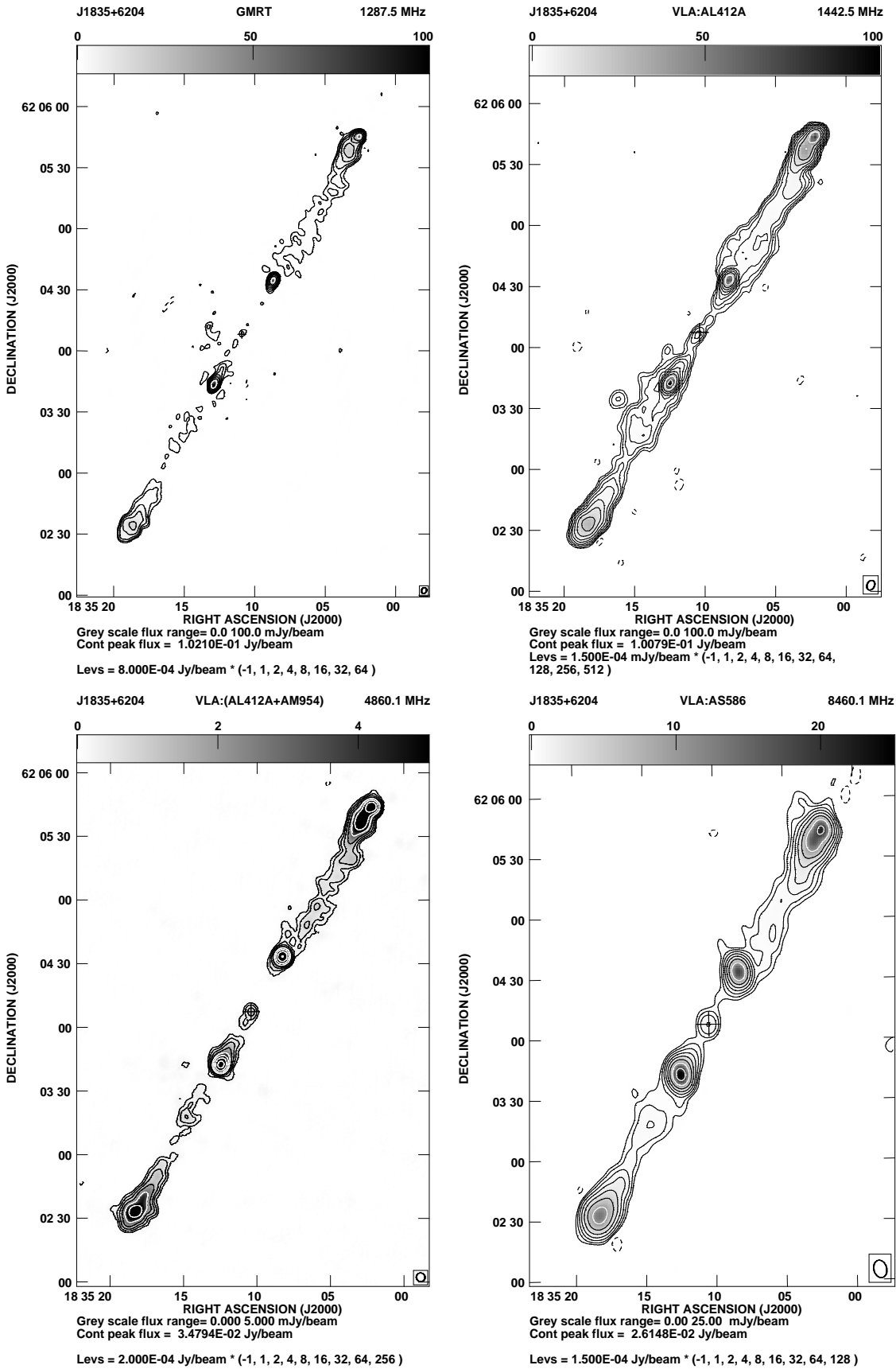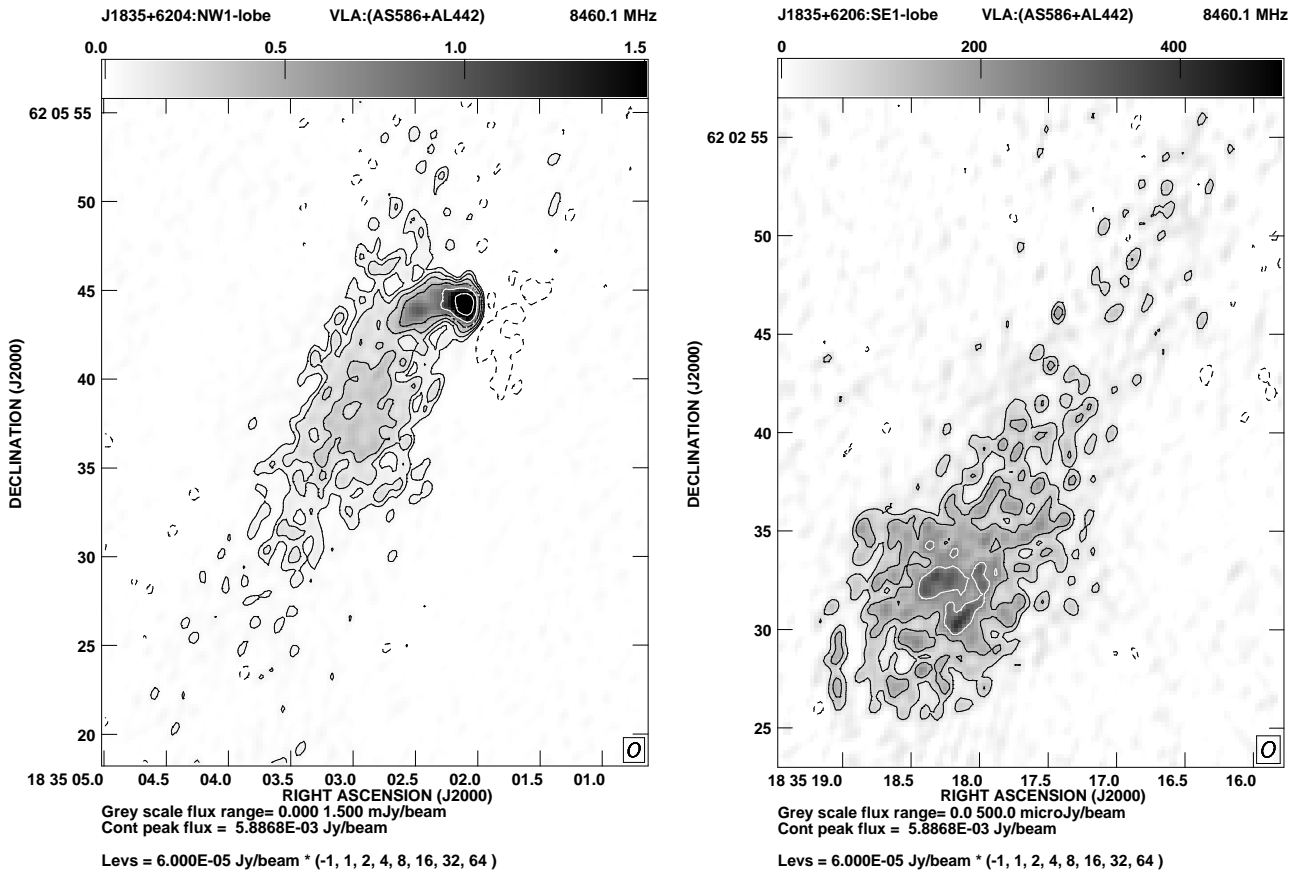
Figure 2 – *continued*

Figure 3. High resolution X band radio images of NW1 and SE1 lobes of J1835+6204 are shown. The images of two lobes of this source is from the same map produced from the data obtained by combining two separate data sets. The resolution of these images is 0.82×0.59 arcsec^2 at $PA \sim 336°$ and the rms noise is 21 μJy/beam. The exact frequency, and the name of the telescope and the observation codes of the data sets used to produce this map are given at the top of each image. The peak flux density, greyscale level, 1st contour and the contour levels are all given at the bottom of each image.

Table 1. Observing log: Column 1 shows the source name; columns 2 and 3 show the name of the telescope, and the array configuration for the VLA observations; column 4 shows the frequency of the observations; while bandwidth and on-source observing time have been listed in column 5 and 6, respectively. The dates of the observations are listed in column 7, and column 8 shows the project code of the data. All GMRT data are our observations.

Source	Teles-cope	Array Conf.	Obs. Freq. (MHz)	Bandwidth used for mapping (MHz)	On-source observing time (min)	Obs. Date	Project code
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
J0041+3224	GMRT		153.28	5.00	360	11-AUG-2006	10CKa01
	GMRT		240.25	5.00	307	17-JUL-2006	10CKa01
	GMRT		331.88	12.50	270	15-JUN-2006	10CKa01
J1835+6204	GMRT		153.25	5.00	240	01-JUN-2006	10CKa01
	GMRT		240.25	5.00	280	14-MAY-2006	10CKa01
	GMRT		332.50	12.50	371	16-JUN-2006	10CKa01
	GMRT		605.81	12.50	280	14-MAY-2006	10CKa01
	GMRT		1287.50	12.50	360	13-JUL-2006	10CKa01
	VLA	B	1442.50	50.00	8.0	18-FEB-1997	AL412A
	VLA	A	1440.50	50.00	251	21-OCT-2000	AS697
	VLA	D	4860.10	100.00	14	27-JUN-2008	AM954
	VLA	B	4860.10	100.00	7.0	18-FEB-1997	AL412A
	VLA	D	8460.10	100.00	23	23-AUG-1996	AS586
	VLA	B	8460.10	100.00	30	07-OCT-1998	AL442

per have been assigned errors determined by propagating errors from the errors of directly measured flux densities.

For both the sources the component designations are as follows. The outer north-western and south-eastern lobes are called NW1 and SE1, respectively. The inner north-western and south-eastern lobes are called NW2 and SE2, respectively. Since in low frequency images the inner doubles appear embedded in the diffuse plasma of the outer cocoons, we could not measure the flux densities of the inner and outer lobes separately on a particular side of the core. In these cases, we measured the total flux densities of the entire NW1 plus NW2 lobes and SE1 plus SE2 lobes, and designated those joint components as the NW1&2 and SE1&2 components respectively. All the inner lobes have edge-brightened structures called hotspots, and hence are of FRII type in morphology. The outer lobes of J0041+3224, as seen in high resolution images, do not have any compact hotspots, but neither do their structures resemble an FRI morphology. In the poor resolution images we see a peak of emission in the outer lobes at both ends of the source along the length. These are most probably the locations of the hotspots when the outer sources were active. These locations can be called warm-spots. Therefore, the outer structures are consistent with the idea that they are created in the previous episode of Jet Forming Activity (JFA).

The case of J1835+6204 is different. At the outer end of the NW1 lobe of J1835+6204, we see a reasonably compact component (hotspot), though the SE1 lobe does not show such a component. However, the radio spectra at the position of the outer hotspots are well described by power laws (see Figure 6). This means that the outer lobes at the location of their hotspots are still fed by jet material (for NW1 lobe) or else the time elapsed since the last jet material reached the outer hotspot is short (for the SE1 lobe). In other words, the spectra have not yet suffered sufficient losses to develop curvature. Both inner and outer pairs of hotspots are more or less co-linear with the core, and both the doubles are also reasonably aligned with each other.

3.2　Notes on the sources

J0041+3224 (B2 0039+32): this DDRG has a very elongated and cylindrical shape. The projected linear sizes of the inner and outer doubles are 171 and 969 kpc respectively. This makes the outer structure a giant radio galaxy. The host galaxy of this source is a 20th magnitude galaxy [NASA/IPAC Extragalactic Database (NED)]. This source does not have a measured redshift; however, we adopt the estimated redshift of 0.45 quoted by Saikia et al. (2006). In low frequency images we can see that the inner double is ploughing its way through the cocoon material of the outer lobes. The inner lobes have compact hotspots **(see Saikia et al. 2006)** which are characteristic of FRII radio galaxies. The outer doubles do not have any hotspots. Radio images at two frequencies are shown in Figure 1

J1835+6204 (B1834+620): this DDRG also has a very elongated and cylindrical shape. The projected linear sizes of the inner and outer doubles are 369 and 1379 kpc, respectively. The outer structure of this radio galaxy is also of giant size. Its host galaxy is a 19.7th magnitude galaxy with a redshift of 0.5194 ± 0.0002 (Schoenmakers et al. 2000a). The inner

double of this DDRG also appears to be ploughing through the cocoon plasma of previous JFA. Unlike J0041+3224, one of the outer lobes (NW1 lobe) of this sources has a compact hotspot-like feature. From Figure 3, it is clear that the jet of NW1 lobe in the previous episode had bent in the western direction by 90 degree with respect to the original jet axis. Figure 2 and 3 display the images of this source at various radio frequencies and resolutions.

3.3　Spectra

The observational parameters and flux densities of the outer and inner lobes/doubles of J0041+3224 and J1835+6204 are listed in Table 2 and 3. Obviously, all the uv data used for imaging have different shortest baselines, and so we re-mapped the fields of both sources at high frequencies with a similar lower uv cutoff to image the inner double free from outer diffuse emission. We tabulate the high-frequency flux densities and their errors **in Table 4** for the lobes and total inner doubles of J0041+3224 and J1835+6204, which were measured from the maps re-made with similar lower uv cutoff. No appreciable curvature is visible in the spectra of the components (except the core) of the inner doubles of both the sources within our observable frequency range. Therefore, we have fitted power laws to the spectral data of the inner lobes and the total inner doubles with the flux density and error values tabulated by Konar et al. (2012, in prep). The flux density measurements of J0041+3224 at 617, 1287, 4860 and 8460 MHz have been made from the FITS maps published by Saikia et al. (2006). A power law has been fitted to only three high frequency points (namely, at 1287, 4860 and 8460 MHz) tabulated **in Table 4** to constrain the spectra of NW2 lobe, SE2 lobe and the total inner double of J0041+3224. These spectra have been extrapolated to lower frequencies. Expressing S in mJy and ν in MHz, the best fitting power laws for the SE2 lobe and NW2 lobe of J0041+3224 are $S_{SE2}(\nu) = (55865.9 \pm 14510) \times \nu^{-(0.698\pm0.031)}$ and $S_{NW2}(\nu) = (45541.8 \pm 10940) \times \nu^{-(0.801\pm0.028)}$ respectively. The best fitting power law for the total inner double is $S_{inn} = (93737.3 \pm 23130) \times \nu^{-(0.724\pm0.030)}$. For J1835+6204, we have fitted power laws to only four high frequency data points (namely, at 1287, 1440, 4860 and 8460 MHz) tabulated **in Table 4** to constrain the spectra of SE2 lobe, NW2 lobe and entire inner double. The best fitting power laws for SE2 and NW2 lobes are $S_{SE2}(\nu) = (81808.3 \pm 4859) \times \nu^{-(0.880\pm0.007)}$ and $S_{NW2}(\nu) = (35066.7 \pm 8009) \times \nu^{-(0.830\pm0.028)}$ respectively. The best fitting power law of the total inner double is $S_{inn}(\nu) = (114830.0 \pm 12800) \times \nu^{-(0.860\pm0.014)}$. For this source as well, we have used the same units of S and ν to express these power laws. The fits of both the sources along with the observed data points are plotted in Figure 4. The flux densities of the integrated source, inner and outer doubles, are given in Table 5 and 6. The integrated spectra of the entire source, and the spectra of the core-subtracted outer and inner doubles are presented in Figure 4. Clearly there is steepening in the spectra of the integrated source and the outer lobes at high frequencies. The steepening of the spectra at higher frequency is presumably due to spectral ageing (due to synchrotron and inverse-Compton losses). Low-frequency measurements from our observations supplemented with those from the literature show significant

Table 2. The observational parameters and flux densities of the outer and inner lobes of J0041+3224. The description of the columns are as follows. Column 1 shows frequency of observation; columns 2, 3 and 4 show the beam size and its orientation; column 5 lists the rms noise of the map; **column 6 presents the integrated flux density**; columns 7, 10, 13, 16 list the component designation. Columns 8, 11, 14 and 17 list the peak flux density of the component. Columns 9, 12, 15 and 18 list the flux density of the component. The component designations are explained in Section 3.1.

Freq. MHz	Beam size ''	''	°	rms mJy /b	S_I mJy	Cp	S_p mJy /b	S_t mJy	Cp	S_p mJy /b	S_t mJy	Cp	S_p mJy /b	S_t mJy	Cp	S_p mJy /b	S_t
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)	(14)	(15)	(16)	(17)	(18)
G153.28	28.14	19.55	349	8.88	7095	SE1&2	1538	4250				NW1&2	1054	2927			
G240.25	16.58	13.66	304	4.11	4374	SE1&2	1020	2615				NW1&2	581	1752			
G331.88	09.46	08.41	61	4.87	4206	SE1&2	960	2465				NW1&2	493	1640			
G616.81	6.50	4.80	165	0.49	2211	SE1&2	567	1289				NW1&2	274	856			
G1287.44	2.61	2.33	25	0.20	1016	SE1	13	232	SE2	231	373	NW1	12	262	NW2	92	145
V1400.00	14.65	12.64	11	0.13	940	SE1&2	338	582				NW1&2	139	360			
V4860.10	4.04	3.65	13	0.03	296	SE1	6	36	SE2	126	154	NW1	7.4	49	NW2	42	52
V8460.10[†]	2.41	2.26	19	0.02	160	SE1	1	9.31	SE2	75	100	NW1	1.8	16	NW2	21	32

[†]: the flux densities are measured after primary beam correction, which was not done in Saikia et al. 2006.

Table 3. The observational parameters and flux densities of the outer and inner lobes of J1835+6204. The column descriptions and the component designations are exactly the same as in Table 2.

Freq. MHz	Beam size ''	''	°	rms mJy /b	S_I mJy	Cp	S_p mJy /b	S_t mJy	Cp	S_p mJy /b	S_t mJy	Cp	S_p mJy /b	S_t mJy	Cp	S_p mJy /b	S_t mJy
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)	(14)	(15)	(16)	(17)	(18)
G153.25	26.18	21.27	331	13.1	5510	SE1&2	811	2539				NW1&2	1263	2898			
G240.25	16.68	13.71	337	7.31	5762†	SE1&2	740	2653				NW1&2	1116	3141			
G332.50	10.28	09.51	21.5	1.72	3375	SE1&2	441	1555				NW1&2	506	1813			
G605.81	8.94	8.40	341	1.25	1813	SE1&2	261	800				NW1&2	257	980			
G1287.50	2.99	2.40	338	0.17	862	SE1&2	102	390				NW1&2	70	462			
V1442.50[a]	4.99	3.66	334	0.02	708	SE1&2	101	318				NW1&2	75	388			
V1440.50[b]	1.35	1.00	357	0.04	657	SE1	2.7	153	SE2	47	135	NW1	23	273	NW2	30	83
V4860.10[a]	1.52	1.12	324	0.02	200	SE1	1.1	37	SE2	17	46	NW1	11	76	NW2	13	33
V4860.10[c]	15.31	10.03	46	0.03	245	SE1&2	44	110				NW1&2	48	133			
V4860.10[d]	1.71	1.14	322	0.02	257	SE1	2.1	70	SE2	21	50	NW1	12	96	NW2	15	32
V8460.10[e]	0.79	0.57	336	0.02	75	SE1	0.20	5.6	SE2	5.5	26	NW1	2.9	26	NW2	3.4	18
V8460.10[f]	8.79	6.45	14.5	0.04	148	SE1&2	26	68				NW1&2	22	78			
V8460.10[g]	0.82	0.59	336	0.02	141	SE1	0.4	34	SE2	5.9	30	NW1	3.1	53	NW2	3.7	18

†: Comparing with the flux density values at all other frequencies, it is likely to have some systemic error. [a]: Project code- AL412A, [b]: Project code- AS697, [c]: Project code- AM954, [d]: Project code- (AL412A+AM954), map from combined data, [e]: Project code- AL442, [f]: Project code- AS586, [g]: Project code- (AS586+AL442A), map from combined data.

Table 4. The flux densities of the inner doubles of J0041+3224 and J1835+6204 from our measurements. The common lower uv cut-off of 3.0 kλ has been applied to the data sets to image the inner double free from outer diffuse emission. The GMRT L-band and VLA high-frequency data have been used to image the inner doubles to constrain the spectra of the components of the inner doubles. The description of the table is as follows: column 1: frequency of observations with the name of the source; columns 2-4: major axis, minor axis and position angle of the synthesized beam; columns 5, 7 and 9: total flux density of inner double (without core), flux density of inner SE lobe and that of inner NW lobe; columns 6, 8 and 10: the errors of the flux densities; column 11: project code of the data used.

Freq. (MHz)	Beam size ''	''	°	$S_{inn}^{tot} - c$ (mJy)	Err (mJy)	S_{SE2} (mJy)	Err (mJy)	S_{NW2} (mJy)	Err (mJy)	Project code
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
J0041+3224										
1287.5	2.61	2.33	25	518	28	373	26	145	10	a
4860.1	4.04	3.65	13	206	8	154	8	52	3	b
8460.1	2.41	2.26	19	132	5	100	5	32	2	b
J1835+6204										
1287.5	2.10	1.84	338	250	12.70	154	10.78	96	6.72	g
1440.5	1.35	1.00	357	218	7.94	136	6.80	82	4.10	h
4860.1	1.67	1.09	322	79	2.84	47	2.35	32	1.60	i
8460.1	0.82	0.59	336	48	1.73	29	1.45	19	0.95	j

a: Target of opportunity observation, b: AS741, g: 10CKa01, h: AS697, i: AL412A+AM954, j: AS586+AL442

Table 5. The integrated flux densities, the total flux densities of the inner and outer doubles, and the uncertainties on these fluxes for J0041+3224. The total flux densities of the outer doubles have been estimated by subtracting the flux densities of the inner doubles from the integrated flux densities. The errors on the total flux densities of the outer doubles are propagated from the errors of the inner double and integrated flux densities. The column designations are self evident.

Freq. MHz	Cp	S_t mJy	Error mJy	Cp	S_t mJy	Error mJy	Cp	S_t mJy	Error mJy	Reference and comment
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
74.00	Int	10882	2176				E1+W1	6727	2464	VLSS
153.28	Int	7095	1064				E1+W1	4402	1279	1
240.25	Int	4374	656				E1+W1	2603	840	1
331.88	Int	4206	631				E1+W1	2804	760	1
408	Int	2200	200							2
616.81	Int	2184	153				E1+W1	1289	319	1
1287.44	Int	1016	71	E2+W2	518*	28[a]	E1+W1	498	76	1
1400.00	Int	940	47							3
1400.00	Int	900	100							4
1400.00	Int	967	48							NVSS
1400.00	Int (avg)	936	40				E1+W1	441	167	5
4830[p]	Int	291	40				E1+W1	89	82	6
4850	Int	269	15							7
4850	Int	273	35							8
4850[q]	Int (avg)	271	19				E1+W1	70	73	9
4840 (Avg)							E1+W1 (Avg)	80	27	10
4860.10	Int	296	15	E2+W2	206*	08[b]	E1+W1	90	17	1
8460.10	Int	160	08	E2+W2	132*	05[b]	E1+W1	28	9.4	1

(avg): The average flux density of the values in the previous rows for the same frequency.

(Avg): The average of [p] and [q]

*: The power law spectra of the total inner lobes (without the core) were constrained by a least-squares fit with these data, and extrapolated to the lowest frequency. The flux density of total inner double (core subtracted) at all other frequencies was estimated from the fitted power law.

[a]: Fluxes of E2 and W2 have been separately measured by us from the FITS image with **the AIPS task 'TVSTAT'**. Errors for E2 and W2 were assumed to be 7 per cent in both cases; hence, the error of on E2+W2 is the quadratic sum of the errors on E2 and W2.

[b]: The same procedure was followed as in [a]; however, the errors for E2 and W2 were assumed to be 5 per cent.

References and comment: The references are to the total flux densities.

VLSS: VLA Low-frequency Sky Survey (Cohen et al. 2007). We have measured the flux from the FITS file of the map. 20 per cent error has been assumed for the 74 MHz integrated flux.

NVSS: NRAO VLA Sky Survey (Condon et al. 1998). 5 per cent error in integrated flux has been assumed.

1: This work. Integrated flux values have been measured by us directly from the maps with **the AIPS task 'TVSTAT'**. 2: Colla et al., 1970. Error has been assumed to be equal to the completeness limit of the survey. 3: Saikia et al., 2006. Flux values are measured by us from the FITS maps . 4: White & Becker, 1992. Error has been assumed to be equal to the completeness limit of the survey. 5: The integrated flux at 1400 MHz has been averaged. The errors in three independent measurements have been propagated to calculate the error in average integrated flux. 6: Griffith et al., 1990. The error has been assumed to be equal to the completeness limit of the survey. 7: Becker, White & Edwards 1991. Error is 5.6 per cent as quoted by NED. 8: Gregory & Condon, 1991. Error is 12.8 per cent as quoted by NED. 9: The integrated flux at 4850 MHz has been averaged. The errors in two independent measurements have been propagated to calculate the error in average integrated flux. 10: Since the outer lobes have errors as large as the fluxes at 4850 and 4830 MHz, we averaged the total outer lobe fluxes at 4850 and 4830 MHz to get a better flux value with less error at some average frequency of 4840 MHz. The errors on the outer total fluxes at 4850 and 4830 MHz have been propagated to calculate the error at the averaged frequency of 4840 MHz. For the spectrum of total outer lobes this averaged point is considered instead of using the 4850 and 4830 MHz points separately.

straightening of the spectra below ∼300 MHz. This low frequency straight part of the spectra can be characterised by a power law whose spectral index most likely represents the injection spectral index (α_{inj}). Further analysis of these issues is described in Section 4.

3.4   Radio core

The J2000.0 positions of the radio core estimated from our high-resolution images are RA: $00^h41^m46\overset{s}{.}11$, DEC: $+32°24'52\overset{''}{.}83$ for J0041+3224 and RA: $18^h35^m10\overset{s}{.}41$, DEC: $+62°04'07\overset{''}{.}46$ for J1835+6204. For J0041+3224, a weak radio core has been detected only at 4860 MHz with a flux density of 0.57 mJy, and a weak peak of emission at 8460 MHz with a brightness of 0.2 mJy/beam at the position of the core has been found by Saikia et al. (2006).

The radio source J1835+6204 has a clear core detection at L band, C band and X band. The core flux densities as measured by us from our observations as well as from the VLA archival data have been compiled in Table 7. Different observations were done at different epochs, enabling us to

Table 6. The integrated flux densities, total flux densities of inner double and that of outer double, and the errors of these flux densities for J1835+6204 are listed in this table. Similar procedures as in Table 5 have been followed to compile the flux densities. Column designations are self evident.

Freq. MHz	Cp	S_t mJy	Error mJy	Cp	S_t mJy	Error mJy	Cp	S_t mJy	Error mJy	Reference and comment
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
38.00	Int	19100	1910							1
38.00	Int	13700	1370							1
38.00	Int(avg)	16400	1175				SE1+NW1	11371	1772	2
153.25	Int	5510	826				SE1+NW1	3994	850	3
240.25	Int	5929	889				SE1+NW1	4899	900	3
332.50	Int	3375	506				SE1+NW1	2596	517	3
326.00	Int	2970	119							4 (WENSS)
605.81	Int	1813	127				SE1+NW1	1348	143	3
1287.50	Int-c	859	60	SE2+NW2	250*	13^p	SE1+NW1	616	70	3
1400.00	Int	800	40				SE1+NW1	574	52	5 (NVSS)
1442.50^a	Int-c	708	35	SE2+NW2	218*	8^q	SE1+NW1	488	48	6
4850.00	Int	200	19							7
4850.00	Int	195	29							8
4850.00	Int(avg)	198	17				SE1+NW1	120	28	9
4860.10^b	Int-c	245	12	SE2+NW2	79*	3^q	SE1+NW1	167	17	10
8460.10^c	Int-c	148	7	SE2+NW2	48*	1.7^q	SE1+NW1	100	11	11

(avg): The average flux density of the previous rows at the same frequency.

∗: Power law spectra of total inner lobes (without the core) were constrained by least-squares fit with these data, and extrapolated to the lowest frequency. Flux density of total inner double (core subtracted) at all other frequencies was estimated from the fitted power law.

^{a}: From the archival VLA data set with project code, AL412A.

^{b}: From the archival VLA data set with project code, AM954.

^{c}: From the archival VLA data set with project code, AS586.

^{p}: Flux densities of SE2 and NW2 were separately measured by us from the FITS image with **the AIPS task 'TVSTAT'**. Errors of SE2 and NW2 have been assumed to be 7 per cent apiece; hence, the error of SE2+NW2 is the quadratic sum of errors of E2 and W2.

^{q}: Same procedure has been followed as in ^{q}; however, the errors for SE2 and NW2 were assumed to be 5 per cent apiece.

Reference and comment: The references are to the total flux densities.

WENSS: Westerbork Northern Sky Survey (Rengelink et al., 1997). 4 per cent error in total flux density has been assumed.

NVSS: NRAO VLA Sky Survey (Condon et al. 1998). 5 per cent error in integrated flux has been assumed.

1: Both Integrated flux density and error values are from 8C survey (Hales et al., 1995) as quoted by Nasa Extragalactic Database (NED). 2: This total flux density is the average of the two values of 38 MHz fluxes as quoted by NED. 3: The integrated flux densities are from our GMRT measurements (project code: 10CKa01). 4: Westerbork Northern Sky Survey (WENSS, Rengelink et al., 1997). 5: NRAO VLA Sky Survey (NVSS, Condon et al. 1998). 6: The integrated flux density is from the data with project code, AL412A and inner double flux density is from the data with project code, AS697. 7: The integrated flux density and error is from Gregory & Condon (1991). 8: The integrated flux density and error is from Becker, White & Edwards (1991). 9: The integrated flux density is the average of the above two rows. The error has been propagated from those in the above two rows. 10: The integrated flux density is from the data with project code, AM954. 11: The integrated flux density is from the data with project code, AS586.

conclude that no prominent variability has been found over a duration of 12 years.

4 SPECTRAL AGEING ANALYSIS

In order to determine the spectral age in different parts of the lobes, i.e. the time elapsed since the radiating particles were last accelerated, we apply the standard theory describing the time-evolution of the emission spectrum from particles with an initial power-law energy distribution characterised by an injection spectral index and distributed isotropically in pitch angle relative to the magnetic field direction (JP model, see Jaffe & Perola, 1973). Our assumption is that the initial energy distribution over the entire frequency/energy range of the spectrum of a blob of relativistic plasma is characterised by the injection spectral index (α_{inj}).

After sufficient amount of time is elapsed, the synchrotron spectrum develops a curvature at high frequencies. This curvature is characterised by a spectral break frequency. The spectral break frequency (ν_{br}) above which the radio spectrum steepens from the injected power law slope is related to the spectral age and the magnetic field strength through

$$\tau_{rad} = 50.3 \frac{B^{1/2}}{B^2 + B_{CMB}^2} \{\nu_{br}(1+z)\}^{-1/2} \mathrm{Myr}, \qquad (1)$$

where $B_{CMB}=0.318(1+z)^2$ is the magnetic field strength equivalent to the cosmic microwave background radiation; B and B_{CMB} are expressed in units of nT, while ν_{br} is in GHz.

Table 7. Flux densities of the radio core of J1835+6204. The values of the core flux were estimated from two-dimensional Gaussian fits from our high resolution images. The peak values from the fit are adopted as the core flux densities. Errors were calculated by quadratically adding the calibration errors and the `JMFIT` (an `AIPS` task) errors. The fiducial values of calibration errors are assumed to be 5 per cent for VLA data and 7 per cent for GMRT data. The column description of the table is as follows. Columns 1, 2 and 3 show the telescope, the project code and the date of observations; column 4 shows the frequency of observations; column 5 lists the resolution of the map from which we measured the core flux density; and lastly, columns 6 and 7 list the core flux density and its error.

Telescope	Project code	Date of obs.	Freq. MHz	Resolution, PA $'' \times ''$, $^{\deg}$	Flux density mJy	Error mJy
(1)	(2)	(3)	(4)	(5)	(6)	(7)
GMRT	10CKa01	13-JUL-2006	1287.5	02.10×01.84, 338	1.66	0.42
VLA-A	AS697	21-OCT-2000	1440.5	01.35×01.00, 357	1.74	0.09
VLA-B	AL412A	18-FEB-1997	1442.5	04.99×03.66, 334	1.62	0.18
VLA-B	AL412A	18-FEB-1997	4860.1	01.52×01.12, 324	1.53	0.09
VLA-D	AM954	27-JUN-2008	4860.1	15.31×10.03, 46	1.51	0.09
VLA-B	AL442.A	07-OCT-1998	8460.1	00.79×00.57, 336	1.05	0.05
VLA-D	AS586.B	23-AUG-1996	8460.1	08.79×06.45, 14	1.25	0.08

4.1 Determination of $\alpha_{\rm inj}$ and $\nu_{\rm br}$ values

There are three free parameters in the spectral ageing models: the normalisation, the injection spectral index ($\alpha_{\rm inj}$) and the break frequency ($\nu_{\rm br}$). Determination of $\alpha_{\rm inj}$ requires very low frequency data points in the synchrotron spectrum. High frequency data points are required for the accurate determination of $\nu_{\rm br}$. The smaller the lower limit of the frequency span, the better the determination of $\alpha_{\rm inj}$. Once $\alpha_{\rm inj}$ is determined accurately, the higher the upper limit of the frequency span, the better the determination of $\nu_{\rm br}$. However, for high-frequency measurements, it is necessary to be cautious about the loss of flux due to central uv holes in the interferometric data, as well as the loss due to sparse uv coverage and the low surface brightness of the lobes at higher frequencies.

Since low frequency flux measurements are very important in the determination of $\alpha_{\rm inj}$, we have compiled very low frequency measurements from the literature and tabulated in Table 5 and 6. While compiling the low frequency flux densities from the literature, we have discarded those which are very discrepant from the curve obtained by simple free-hand drawing. A simple estimate by eye clearly identifies data that are highly discrepant. Since most of the measurements from the literature are total flux density values, and no flux density values exist for a single lobe component of a source, we have constrained the total synchrotron spectrum of the two outer lobes together for each source to very low frequencies. The flux densities and their errors compiled in Tables 5 and 6 have been used in fitting the spectra of various components of J0041+3224 and J1835+6204. The flux densities of the total inner doubles of both the DDRGs are fitted well with power laws and these best fitting power laws are given in Section 3.3. We have estimated the flux densities of the inner double from the best-fitting power law models at those frequencies for which we could not directly measure the flux densities of the inner doubles due to contamination of diffuse emission of outer double and/or poor resolution. For those estimated flux densities the errors have been calculated using the formula $\sigma_{S_\nu} = \sqrt{\{(\frac{S_\nu}{S_0})^2\sigma_{S_0}^2 + (S_\nu ln\nu)^2\sigma_\alpha^2\}}$, where S_0 (the normalisation), σ_{S_0} (the error in normalisa-

tion) and S_ν are in mJy, ν is in MHz and σ_α is the error in spectral index of a general power law spectrum of the form $\frac{S_\nu}{mJy} = S_0(\frac{\nu}{MHz})^\alpha$.

In order to determine a value of $\alpha_{\rm inj}$, we first fitted the JP (Jaffe & Perola 1973) and Continuous Injection (CI) models (Pacholczyk, 1970) of radiative losses to the flux densities of the outer lobes. We found that the CI model does not provide a good fit to the data. The application of the CI model to the outer structure is physically unjustifiable for DDRGs in general, because a) the jets no longer feed the outer lobes, and b) the outer lobes are so old that they have suffered synchrotron and inverse-Compton losses to a large extent (and possibly adiabatic loss to some extent). So high energy electrons are no longer being replenished in the outer lobes. Therefore, it is reasonable that the CI model does not provide a good fit for the total spectrum of the outer lobes, and instead the JP model is more appropriate. However, for J1835+6204, the NW1 lobe has a clear compact hotspot (a signature of the ongoing feeding by the last ejected jet material in the previous episode of JFA) and both the NW1 and SE1 lobes have single power-law spectra within our observable range of radio frequencies. Even in that case the CI model does not yield a good fit, while the JP model does. This might have something to do with the fact that most of the parts of the lobes except the hotspot are old and the high energy particles are not replenished. Since we cannot constrain the radio spectra of each lobe separately up to very low frequencies, we have assumed that $\alpha_{\rm inj}$ is the same for both the outer lobes, and that $\alpha_{\rm inj}$ remains constant over the entire active phase of JFA.

The various best fitting parameters of the spectra of the outer lobes of our 2 DDRGs are shown in Figure 4 These fits have been used to constrain the values of $\alpha_{\rm inj}$ for the ageing analysis of the outer lobes of the sources. Since there is no visual indication of curvature in the spectra of the inner doubles, we have fitted a pure power law to the inner lobes of both the DDRGs. Then, the total-intensity maps made by us at multiple frequencies were convolved with a common angular resolution (circular beams of $9.''50$ for J0041+3224 and $16.''70$ for J1835+6204). Each lobe was then split into a number of strips, separated approximately by the common
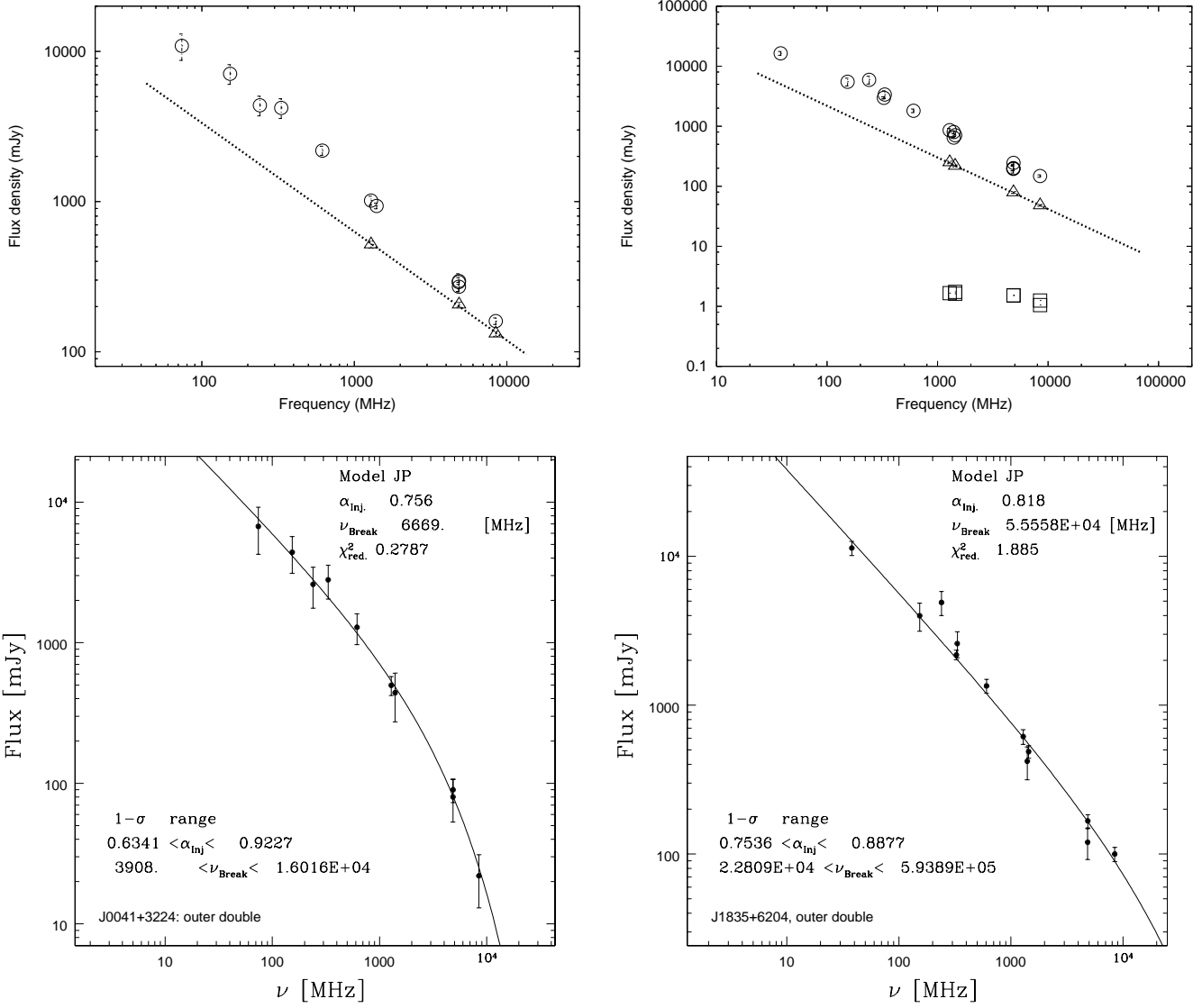
Figure 4. Left upper panel: Integrated and inner double spectra of the source J0041+3224 have been plotted. Open triangles are the inner double flux densities (core subtracted) measured from the high resolution images and open circles are integrated flux densities. The inner double flux densities are fitted with a power law (dashed line) and extrapolated all the way back to 20 MHz to enable the reader to compare its strength with the integrated source. Left bottom panel: Filled circles are the flux densities of the entire outer double. The continuous curve represents the spectral ageing JP model fit. Right upper and right bottom panels are the same as the Left upper and Left bottom panels, respectively, but for the source J1835+6204. The open squares in the upper right panel are the core flux densities (see Table 7).

resolution element (with which the maps were convolved) along the axis of the source, and the spectrum of each strip has been constrained by fitting a spectral ageing model with our measured flux densities. For fitting the synchrotron ageing model to the observed spectra of our sources, we have used the SYNAGE software (Murgia 1996), which is specifically written for this purpose. SYNAGE was used to fit a JP model to the spectrum of each strip of the outer lobes of our sample sources. For a given source, we fixed the value of α_{inj}, as determined by our fits, shown in Figure 4, for all strips. The fitted spectra of some of the strips are shown in Figure 5 and 6. If we treat α_{inj} as a free parameter, then the resulting best-fitting values of α_{inj} for different strips show

significant variation with large error bars. We believe that the large errors of α_{inj} in the fit are possibly due to the large uncertainty in the surface brightness for each strip. So, for all the strips of both lobes of each source, we have fixed the value of α_{inj} to the best fitting value obtained from the JP model fit to the observed spectrum of entire outer double. The values of α_{inj} used for all the strips (of a source) are quoted in the caption of Table 8 and 9 for J0041+3224 and J1835+6204 respectively. Here our assumption is that the jet injected all the relativistic plasma with the same value of α_{inj} over the entire period of the active phase. We would need much more sensitive data to treat α_{inj} as free parameter for the spectrum of each strip. The values of ν_{br} including
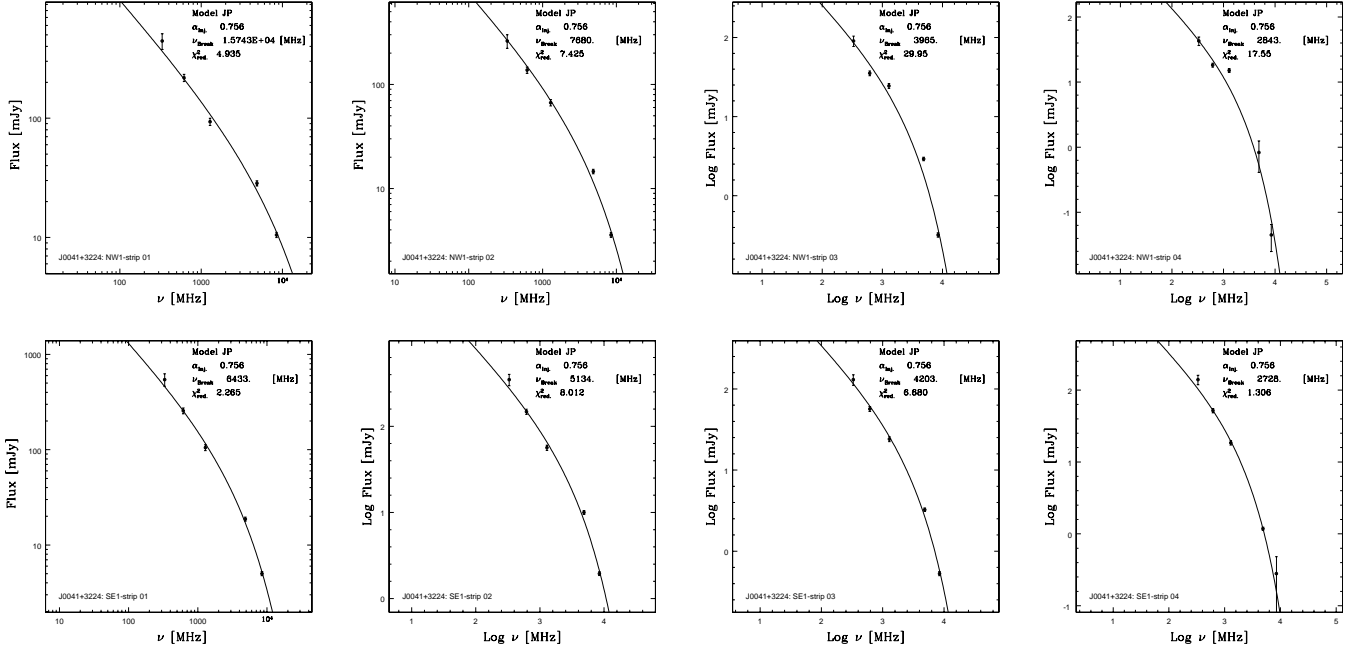
Figure 5. Typical spectra of a few slices for the eastern (upper panel) and western (lower panel) lobes of the outer double of J0041+3224.
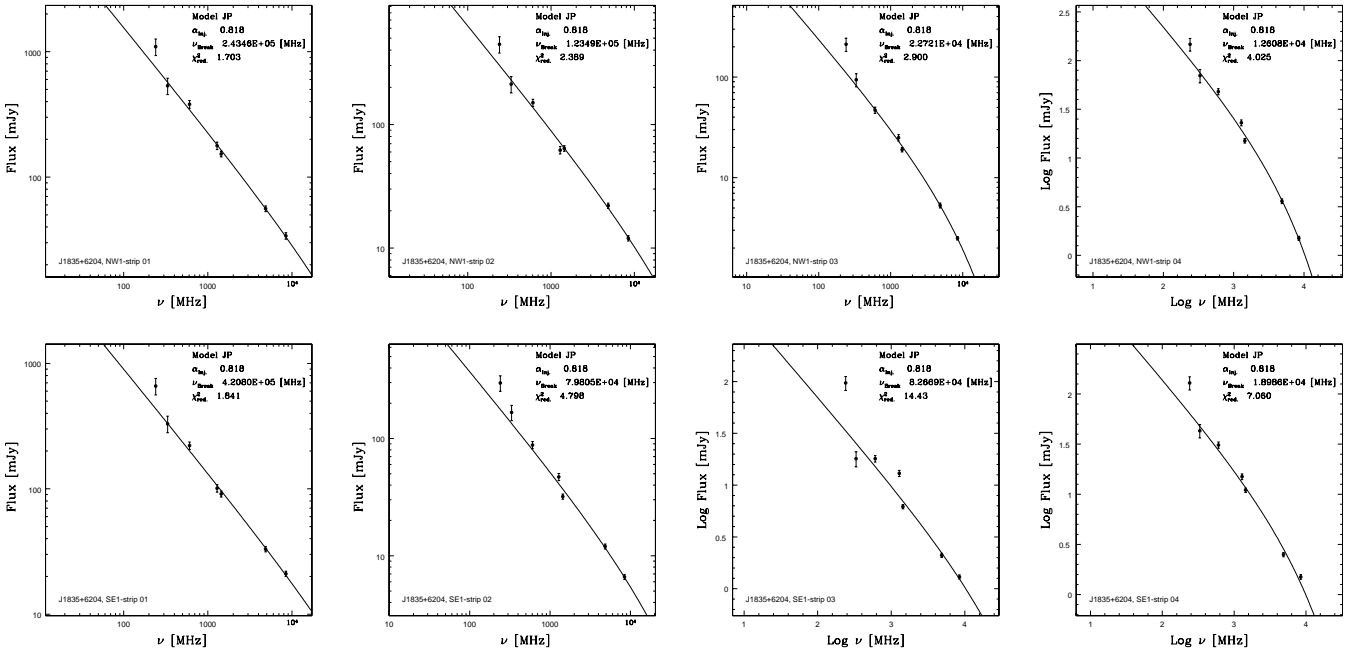

Figure 6. Typical spectra of a few slices for the north-western (upper panel) and south-eastern (lower panel) lobes of the outer double of J1835+6204.

the 1σ errors for the strips of outer lobes are listed in Table 8 and 9. In our previous work, whenever we could constrain α_{inj} for each outer lobe separately, we found that they are not very different (Jamrozy et al. 2008, Machalski et al., 2010), and so there is no strong evidence for significantly different values of this parameter in two opposite lobes of a radio galaxy.

4.2 Magnetic field determination and radiative ages

To determine the spectral age of the particles in each strip, we estimated the magnetic field strength of each strip. The values of the minimum energy density and the corresponding magnetic field, B_{min}, were calculated using the formalism described in Appendix of Konar et al. (2008). We then used

the values of the minimum energy magnetic field and the break frequency (from Table 8 and 9) to calculate the spectral ages of the strips via Equation 1 . While performing spectral ageing analysis we have assumed that 1) the energy losses of the relativistic cocoon plasma are only due to synchrotron and IC processes, 2) the blob of plasma in each strip was injected over a short period compared to the age of the radio galaxy, so that the blob of plasma can be assumed to have been injected in a single shot, 3) there is no mixing of plasma between any two adjacent strips, 4) the magnetic field is in equipartition (the result of Croston et al., 2005 is in favour of this assumption) and does not change appreciably over the lifetime of the radio source and 5) the magnetic field is completely tangled. The spectral ages of the strips as a function of distance are plotted in Figure 7. As expected, the synchrotron ages for both the outer lobes increase with distance from the edges (warm-spots) of the lobes. We have fitted a polynomial to every age-distance plot to extrapolate the curve to the position of the core. The value of the polynomial at the position of the core gives the expected spectral age of the outer lobes. Since in both of our DDRGs there is diffuse relativistic plasma from the outer lobes all the way back to the core, it makes sense to determine the spectral age by this extrapolation method. We could not determine the age of the outer lobe plasma near the core region due to (i) the presence of inner lobes and (ii) the non detection of diffuse plasma in our higher frequency images. While interpreting these numbers, caveats related to the evolution of the local magnetic field in the lobes need to be borne in mind (e.g. Rudnick, Katz-Stone & Anderson 1994; Jones, Ryu & Engel 1999; Blundell & Rawlings 2000). While Kaiser (2000) suggested that spectral and dynamical ages are comparable if bulk backflow and both radiative and adiabatic losses are taken into account in a self-consistent manner, Blundell & Rawlings (2000) argue that this may be so only in young sources with ages much less than 10 Myr. In a study of the FRII type giant radio galaxy, J1343+3758, Jamrozy et al. (2005) find the dynamical age to be approximately 4 times the maximum synchrotron age of the emitting particles.

5   DISCUSSION

We find that the spectral ages of the two different outer lobes of the DDRGs are different. This does not necessarily mean that the two jets started at different points in time. This result may be due to additional loss processes other than synchrotron and IC radiative losses, (e.g., adiabatic expansion loss) and any kind of re-acceleration of the particles that may be at work with differing efficiency in the two lobes. The differing spectral ages then provide information about the asymmetry of the physical conditions (internal and external) of the lobes on opposite sides. A source can have asymmetric jet powers, an asymmetric environment or both, in addition to radiative losses and acceleration of the particles in the lobes.

In the western outer lobe of J0041+3224, we find that the variation of spectral age with distance from the core is not smooth. As we go from the western hotspot towards the core we see that the spectral age becomes almost constant from a distance of 400 kpc all the way to 175 kpc from the core. This may be due to particle re-acceleration (of un-
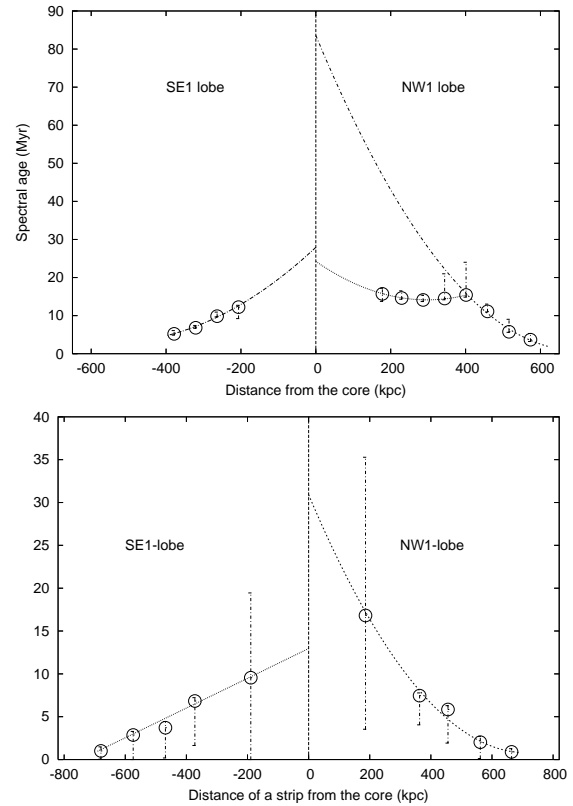

Figure 7. Radiative age of the relativistic particles in the outer lobes plotted against the distance from the radio core using the minimum energy magnetic fields. Top panel: Of the source J0041+3224. Bottom panel: Of the source J1835+6204.

known origin). The outer lobes of J0041+3224 are highly asymmetric which may be due to genuine asymmetry in the density of the ambient medium combined with asymmetric jet power and/or any other asymmetric parameters. Since the outer lobes are not active, the lobe heads are likely to advance with a speed slower than when they were active. Since the radio spectra of the inner doubles are a very close approximation to a single power law, we cannot determine spectral ages for them. However, we have determined upper limits to the spectral ages of the inner doubles with the assumption that the break frequencies are greater than the highest observed frequencies (see Figure 4). We also determined a lower limit to their age by assuming the jet-head advancement speed to be 0.5 times the velocity of light.

As we have discussed, we find different radiative ages for the two lobes of both sources. The actual ages of the lobes are very unlikely to be different as it is not physically plausible for the two jets of an AGN to start at different times. Therefore, our interpretation of the apparent age difference is that processes other than radiative losses, e.g., re-acceleration and adiabatic expansion, are at work. Their relative impact on the energy loss in different lobes might be different, giving rise to different radiative ages. We can therefore obtain a more meaningful spectral age for the outer doubles by taking an average of the spectral ages of the two lobes, provided they are not different by huge factor (e.g.,$\geq$5 times). For J0041+3224, the spectral ages obtained by extrapolating the spectral age vs. distance plot are $\sim$24

and 28 Myr for the NW1 and SE1 lobes respectively, with an average of 26 Myr. If we assume that the re-acceleration of particles took place in the inner parts of the NW1 lobe, then the extrapolated age of this lobe would be ~84 Myr. In such a situation, we must invoke re-acceleration in the entire SE1 lobe as it has a much lower spectral age than the NW1 lobe. It is difficult to infer the true evolutionary history of this source. For J1835+6204, the extrapolated ages are ~31 and 13 Myr for the NW1 and SE1 lobes, respectively, with an average of 22 Myr. The average spectral age of the two warm-spots (t_{ws}) is ~4.44 Myr for J0041+3224 and ~1 Myr for J1835+6204. The spectra of the inner doubles show no curvature, and therefore we could not constrain the break frequency or determine a spectral age. However, the highest frequencies in the power-law spectra of both the DDRGs are 8.46 GHz, which can be assumed to be a lower limit on the break frequency. This leads to upper limits on the spectral ages of the inner doubles of ~5.30 and 3.94 Myr for J0041+3224 and J1835+6204, respectively.

It is known from various observations that the inner hotspots travel with a velocity $\sim 0.1c - 0.5c$ with respect to the host galaxy (Konar et al., 2006; Schoenmakers et al., 2000b; Safouris et al., 2008). So we have estimated the lower limits on the ages (kinematic ages) of the inner doubles for a hotspot velocity of $\sim 0.5c$, which are ~0.34 and 1.34 Myr for J0041+3224 and J1835+6204, respectively. Of course, if one interprets the inner lobes to the bow shock driven by the jets (Brocksopp et al., 2007; 2011; Safouris et al., 2008) then the jet head velocity is supposed to be close to the jet bulk velocity which will be higher than 0.5c. In that situation the lower limit of the ages of the inner doubles will be even lower. The spectral age of the outer hotspot is the time elapsed since the last jet material reached the hotspot in the previous episode of JFA; similarly the age of the inner double is the time since the current episode of JFA started. So the durations of the active phases (t_{activ}) and quiescent phase (t_{quies}) are given by

$$t_{activ} = t_{outd} - (t_{jet} + t_{ws}), \qquad (2)$$

and

$$t_{quies} = (t_{ws} + t_{jet}) - t_{innd} \qquad (3)$$

respectively; where, t_{jet} is the time which any blob of jet material took to travel from the core to the warm-spot in the previous episode, t_{outd} is the age of the outer double, and t_{innd} is the age of the inner double. If we assume that the bulk velocity of the jet fluid is the velocity of light, then the values of t_{jet} (when averaged over the two outer lobes) are 1.58 and 2.25 Myr for J0041+3224 and J1835+6204 respectively. So, the duration of active phase in the previous episode of JFA of J0041+3224 is ~20 Myr (using Equation 2). Since the estimated age of the inner double of J0041+3224 is given by $0.34 \lesssim t_{innd} \lesssim 5.30$ Myr and $t_{jet} + t_{ws} = 6.02$ Myr, the duration of the quiescent phase for this source is given by $0.72 \lesssim t_{quies} \lesssim 5.68$ Myr (using Equation 3). So, quiescent phase of J0041+3224 lies between 3.6 to 28.4 per cent of the active phase of the previous episode of JFA. For J1835+6204 the situation is different. The break frequency is higher than the conventional fiducial value of the highest frequency (~ 100 GHz) up to which the spectrum exists (see Table 9). Therefore, it is unlikely that there is any real curvature in the spectrum. However, SYNAGE tries

Table 8. Results of JP model calculations for J0041+3224 with $\alpha_{\rm inj} = 0.756^{+0.167}_{-0.122}$. Column 1: identification of the strip; column 2: the projected distance of the strip-centre from the radio core; column 3: the break frequency in GHz; column 4: the reduced χ^2 value of the fit; column 5: the estimated magnetic field in nT; column 6: the resulting synchrotron age of the particles in the given strip.

Strip	Dist. kpc	$\nu_{\rm br}$ GHz	χ^2_{red}	$B_{\rm min}$ nT	$\tau_{\rm rad}$ Myr
		NW1-lobe			
NW1-01	573.3	$3.68^{+1.62}_{-0.46}$	4.94	1.86 ± 0.18	$3.68^{+1.62}_{-0.46}$
NW1-02	516.0	$5.76^{+3.27}_{-0.15}$	7.42	1.73 ± 0.17	$5.76^{+3.27}_{-0.15}$
NW1-03	458.6	$11.11^{+1.91}_{-0.17}$	29.95	1.31 ± 0.13	$11.11^{+1.91}_{-0.17}$
NW1-04	401.3	$15.42^{+8.60}_{-0.57}$	17.55	1.12 ± 0.11	$15.42^{+8.60}_{-0.57}$
NW1-05	344.0	$14.50^{+6.49}_{-0.34}$	30.75	1.14 ± 0.11	$14.50^{+6.49}_{-0.34}$
NW1-06	286.6	$14.11^{+1.28}_{-0.42}$	1.70	1.24 ± 0.12	$14.11^{+1.28}_{-0.42}$
NW1-07	229.3	$14.61^{+1.90}_{-0.31}$	14.06	1.16 ± 0.11	$14.61^{+1.90}_{-0.31}$
NW1-08	177.7	$15.72^{+1.45}_{-1.98}$	9.10	1.07 ± 0.11	$15.72^{+1.45}_{-1.98}$
		SE1-lobe			
SE1-01	378.4	$5.20^{+0.87}_{-0.34}$	2.26	2.01 ± 0.20	$5.20^{+0.87}_{-0.36}$
SE1-02	321.0	$6.81^{+0.61}_{-0.15}$	8.01	1.78 ± 0.18	$6.80^{+0.62}_{-0.14}$
SE1-03	263.7	$9.86^{+1.07}_{-0.17}$	6.68	1.42 ± 0.14	$9.86^{+1.07}_{-0.17}$
SE1-04	206.4	$12.23^{+0.47}_{-3.00}$	1.31	1.42 ± 0.14	$12.23^{+0.47}_{-3.00}$

Table 9. Results of JP model calculations for J1835+6204 with $\alpha_{\rm inj} = 0.818^{+0.070}_{-0.064}$. Column 1: identification of the strip; column 2: the projected distance of the strip-centre from the radio core; column 3: the break frequency in GHz; column 4: the reduced χ^2 value of the fit; column 5: the estimated magnetic field in nT; column 6: the resulting synchrotron age of the particles in the given strip.

Strip	Dist. kpc	$\nu_{\rm br}$ GHz	χ^2_{red}	$B_{\rm min}$ nT	$\tau_{\rm rad}$ Myr
		NW1-lobe			
NW1-01	664.2	$243.5^{+1.34\times10^5}_{-121.5}$	1.70	1.85 ± 0.18	$0.90^{+0.37}_{-0.86}$
NW1-02	561.7	$123.5^{+3.94^4}_{-54.7}$	2.39	1.21 ± 0.12	$2.02^{+0.69}_{-1.91}$
NW1-03	456.2	$22.7^{186.3}_{-3.1}$	2.90	0.94 ± 0.09	$5.85^{+0.44}_{-3.92}$
NW1-04	363.1	$12.6^{+30.1}_{-1.0}$	4.02	1.00 ± 0.10	$7.64^{+0.14}_{-3.59}$
NW1-05	186.2	$2.2^{+47.4}_{-1.7}$	0.38	1.08 ± 0.11	$16.82^{+18.46}_{-13.28}$
		SE1-lobe			
SE1-01	679.7	$421.0^{+1.96\times10^5}_{-225.0}$	1.84	1.31 ± 0.13	$1.01^{+0.47}_{-0.96}$
SE1-02	574.2	$80.0^{+1.68\times10^{11}}_{-19.0}$	4.80	1.04 ± 0.10	$2.87^{+0.42}_{-2.87}$
SE1-03	468.6	$82.6^{+3.14\times10^4}_{-24.6}$	14.43	0.67 ± 0.07	$3.72^{+0.72}_{-3.53}$
SE1-04	372.4	$19.0^{+308}_{-2.0}$	7.06	0.85 ± 0.08	$6.83^{+0.39}_{-5.18}$
SE1-05	189.3	$9.58^{+1.81\times10^{11}}_{-7.27}$	0.99	0.86 ± 0.08	$9.54^{+9.90}_{-2.59}$

its best to fit with a curvature. There is no appreciable curvature in the spectra of the outer hotspots (NW1-strip 01 and SE1-strip 01) of J1835+6204 (see Figure 6). Though we have formally obtained an age of ~1 Myr, ageing has not really started at the outer hotspots, as the last ejected jet material is still feeding them, which is particularly clear in the NW1 hotspot. Therefore, we will assume the ages of the outer hotspots to be zero in the sense that they are still fed, and as a result the high energy particles are replenished. In high resolution X band image (Figure 3), we can easily see

that there is a hotspot like feature in NW1 lobe. However we don't see such prominent hotspot like feature in the SE1 lobe. Hotspot formation is a complicated phenomenon. It certainly seems possible that a diffuse hotspot can be formed without the jet having been turned off completely, and that is evident in the radio image of high resolution hotspots in some of the FR II sources published by Leahy et al., (1997). Moreover, if there is a weak compact hotspot like the one in north western lobe of 3C105 (see Leahy et al., 1997), we might not be able to see because of the sensitivity of our high resolution X band image in Figure 3. So, it is difficult to conclude whether the SE1 lobe is still fed by the jets or not. Here we argue from both points of view.

First we assume that NW1 and SE1 lobes are still fed by the jets of previous episode. The last jet material has been ejected from the nucleus, but the jet material is still travelling down the jets. However, we do not know how long ago the last jet material was ejected from the nucleus. We assume that t'_{jet}, which is $\lesssim t_{jet} (= 2.25$ Myr$)$, is the time since the last jet material was ejected. So to calculate t_{activ} and t_{quies}, t_{jet} and t_{ws} should be replaced by t'_{jet} and 0 in Equations 2 and 3 respectively. However, t'_{jet} must be greater than the lower limit of t_{innd} as the last jet material of previous cycle had to be ejected before the current episode started. So $t'_{jet} \gtrsim 1.34$. Again, if the last ejected jet material has already reached the hotspots, then $t'_{jet} = t_{jet} = 2.25$ Myr. So, we ultimately get $1.34 \lesssim t'_{jet} \lesssim 2.25$ Myr. The upper limit of the age of the inner double can be at most t_{jet}, not greater than that. So, $t_{innd} \leq t_{jet}$. Therefore, we get $1.34 \lesssim t_{innd} \lesssim 2.25$ Myr. Considering that t_{quies} must be a positive number, and the limits of t_{innd} and t'_{jet}, we get $19.75 \lesssim t_{activ} \lesssim 20.66$ Myr and $0 \lesssim t_{quies} \lesssim 0.91$ Myr. The limits of t_{activ} are so close to each other that the average of those limits is a good approximation of the t_{activ} which is ~20.20 Myr. So, in this range of quiescent phase, the last ejected jet material can travel from the nucleus up to 279 kpc, and not more than that. So the quiescent phase of J1835+6204 has lasted less than ~ 4.5 per cent of the active phase of the previous episode of JFA.

If we assume that the lack of a hotspot in SE1 lobe is real, then it may suggest that jet activity has ceased there and we expect that the hotspots of the inner double are the current jet termination points. In such a case, the presence of a compact hotspot in the NW1 lobe is the feature that needs to be explained. Since hotspots of the inner double are the current jet termination points, then we would expect the hotspot in the NW1 lobe to disappear at some point. Assuming also that the jet turns off on both sides at the same time, the only way we can see a hotspot in the NW1 lobe and not in the SE1 lobe is if the NW1 lobe is on the far side of the source, so that we are seeing it at an earlier time. For the sake of a simple calculation we assume that the source is symmetrical. Let the true lobe length (from the core to the hotspot) be L and the angle to the line of sight be θ. Then $L_{obs} = 2L\sin\theta$ and the distance along the line of sight between the hotspots is given by $D_{LTT} = 2L\cos\theta$, where L_{obs} is the observed lobe length and D_{LTT} is basically the Light Travel Time (LTT) multiplied by speed of light (c). So D_{LTT} is $L_{obs}\cot\theta$. If we require $\theta > 45$ degrees (unification arguments) then the maximum value of $D_{LTT} = L_{obs} \sim 1.4$ Mpc. So the maximum time difference between the hotspots

is 4.6 Myr. Suppose the jet was disconnected from the lobes at a time t'_{jet} in the past and the effects travel down the old jet channel at a speed $\beta_{jet}c$. If we see no hotspot in the SE lobe, then we know that $\beta_{jet}ct'_{jet} > L$. But if we see a hotspot in the NW1 lobe, which we see at an earlier time, then we know $\beta_{jet}c(t'_{jet} - D_{LTT}/c) < L$, or $\beta_{jet}ct'_{jet} < (L + \beta_{jet}D_{LTT})$. For simplicity if we assume $\theta = 45$ degrees and $\beta_{jet} = 1$, then we have $ct'_{jet} \gtrsim 1$ Mpc and $ct'_{jet} \lesssim (1 + 1.4)$ Mpc. So, essentially we obtain $3.3 \lesssim t'_{jet} \lesssim 7.9$ Myr. So the jet switch-off time has to be more than 3.3 Myr ago and less than 7.9 Myr ago (measured with respect to the observations of the SE1 hotspot). We have already found out that the jet travel time t_{jet} of the last ejected material for J1835+6204 is 2.25 Myr which is less than t'_{jet}. Therefore, we can calculate the hotspot fading away time ($t_{hs,fade}$) from this analysis. The hotspot fading time can be given by

$$t_{hs,fade} = t_{ws} = t'_{jet} - t_{jet}. \tag{4}$$

Considering the limits of t'_{jet} for the source J1835+6204, we obtain $1.05 \lesssim t_{hs,fade} \lesssim 5.65$ Myr. Even if the hotspot fade away time is the lower limit of this range, i.e., 1.05 Myr which is enough for the SE1 hotspot to get diffuse, as it takes only sound crossing time which is $\sim 10^5$ yr (Kaiser et al., 2000). From Equation 4, the age of the warm spot (t_{ws}) in SE1 lobe can be given by $1.05 \lesssim t_{ws} \lesssim 5.65$ Myr. Using the magnetic field value of 1.31 ± 0.13 nT of the SE1-strip01 from Table 9 into Equation 1, we obtain $13.5 \lesssim \nu_{br} \lesssim 390$ GHz. So, our estimation of warmspot age in the SE1 lobe of J1834+6204 in our interpretation of absence of the SE1 hotspot is fully consistent with the fact that we don't see any curvature in the spectrum of SE1-strip01 (see Figure 6) in our observable range of frequencies. This particular set of numbers doesn't seem to make it so implausible that we are seeing a genuine one-sided hotspot, though these are the most favourable assumptions. As θ approaches 90 degrees, we have to be observing the source at more and more special times to see a hotspot in one lobe and not in the other, since our constraints are basically $1 < \frac{ct_{jet}}{L} < 1 + 2\cos\theta$. So, for any source with roughly symmetrical lobes there will be some time window where we expect to see a hotspot in one lobe but not in the other. Now again in Equations 2 and 3, we replace t_{jet} and t_{ws} by t'_{jet} and 0 (as the spectrum is straight) respectively. Keeping in mind the limits of t'_{jet} and t_{innd}, we obtain from Equations 2 and 3 that $14.1 \lesssim t_{activ} \lesssim 18.7$ Myr and $1.05 \lesssim t_{quies} \lesssim 6.56$ Myr. So, if this interpretation is correct, then the quiescent phase of J1835+6204 can be given by $5.6 \lesssim t_{quies} \lesssim 46.5$ percent of the active phase of the previous cycle.

6 CONCLUDING REMARKS

We have presented the results of multifrequency radio observations of two DDRGs, J0041+3224 and J1835+6204, using both the GMRT and the VLA. We have carried out detailed spectral ageing analysis of these two sources, and have come to the following conclusions.

(i) From our observations, it is quite clear that the inner doubles are propagating through the cocoon material deposited in the previous episode of JFA. This may be the

case for most of the DDRGs, even though the old cocoon material is not always visible in images of limited sensitivity.

(ii) Our spectral ageing analysis reveals the youth of the inner doubles and the old age of the outer doubles, consistent with the model in which the central engine is capable of restarting a previously quenched jet of the central engines, and with all previous work.

(iii) The images show that in both cases the pairs of doubles are not significantly misaligned. This suggests that there has not been any realignment of the galaxy gas disk, accretion disk and/or jet axis of the central engine.

(iv) The radio core of J1835+6204 shows no evidence of variability, whereas Konar et al. (2006) and Jamrozy et al. (2007) have reported high variability in the cores of the restarting radio galaxies J1453+3308 and 4C29.30 respectively. The results we present here suggest that strong core variability may or may not be observed in restarting radio galaxies.

(v) The synchrotron age of the emitting particles in both the lobes of the outer doubles increases with distance from the edges of the lobes, as expected from the backflow model of FRII radio galaxies (**This behaviour is also seen in lobed FR I sources. See Parma et al., 1999 and Laing et al., 2011.**). There appears to be either re-acceleration of particles in the inner part of the NW1 lobe of J0041+3224, or adiabatic loss in the middle of this lobe. The reasons for the re-acceleration or adiabatic loss are not known. Further investigation is required to understand this phenomenon.

(vi) The duration of the quiescent phase of J0041+3224 is greater than 0.72 Myr but less than 5.68 Myr, which is between $\sim 4-28$ per cent of the duration of the active phase of the previous episode of JFA, provided the age of the outer double is 26 Myr. If the age of the outer double is 84 Myr, then the duration of quiescent phase is $\sim 0.9-7$ per cent of the duration of the active phase of the previous episode of JFA. The duration of the quiescent phase of J1835+6204 is less than $\sim$1 Myr, which is ~ 4.5 per cent of the active phase of the previous episode of JFA. It is therefore clear that the quiescent phase can be as small as a few per cent of the active phase duration.

(vii) Our age estimates of J1835+6204 are fully consistent with the outer lobes being still fed by the last ejected jet material, though the alternative explanation of one sided visible hotspot because of the LTT effect is also physically plausible.

REFERENCES

Baars J.W.M., Genzel R., Pauliny-Toth I.I.K., Witzel A. 1977, A&A, 61, 99

Barnes J.E., Hernquist L., 1996, ApJ, 471, 115

Becker R.H., White R.L. and Edwards A.L., 1991, ApJS, 75, 1

Belsole E., Worrall D.M., Hardcastle M.J., Croston J.H., 2007, MNRAS, 381, 1109

Blundell K.M., Rawlings S., 2000, AJ, 119, 1111

Brocksopp, C., Kaiser, C.R., Schoenmakers, A.P., de Bruyn, A.G., 2007, MNRAS, 382, 1019

Brocksopp, C., Kaiser, C.R., Schoenmakers, A.P., de Bruyn, A.G., 2011, MNRAS, 410, 484

Cohen A.S., Lane W.M., Cotton W.D., Kassim N.E., Lazio T.J.W., Perley R.A., Condon J.J., Erickson W.C, 2007, AJ, 134, 1245

Colla G. et al., 1970, A&AS, 1, 281

Condon J.J., Cotton W.D., Greisen E.W., Yin Q.F., Perley R.A., Taylor G.B., & Broderick J.J. 1998, AJ, 115, 1693

Croston J.H., Birkinshaw M., Hardcastle M.J., Worrall D.M., 2004, MNRAS, 353, 879

Croston J.H., Hardcastle M.J., Birkinshaw M., Worrall D.M., 2003, MNRAS, 346, 1041

Croston J.H., Hardcastle M. J., Birkinshaw M., Worrall D.M., Laing R.A., 2008, MNRAS, 386, 1709

Croston J.H., Hardcastle M.J., Harris D.E., Belsole E., Birkinshaw M., Worrall D.M., 2005, ApJ, 626, 733

Evans D.A., Hardcastle M.J., Croston J.H., Worrall D.M., Birkinshaw M., 2005, MNRAS, 359, 363

Gregory P.C. and Condon J.J., 1991, ApJS, 75, 1011

Griffith Mark, Langston G., Heflin M., Conner S., Lehar J., Burke B., 1990, ApJS, 74, 129

Hales S.E.G., Waldram E.M., Rees N., Warner P.J., 1995, MNRAS, 274, 447

Hardcastle M.J., and Worrall D.M., 1999, MNRAS, 309, 969

Hardcastle M.J., and Worrall D.M., 2000, MNRAS, 319, 562

Hota A., Sirothia S.K., Ohyama Y., Konar C., Kim S., Rey S.-C., Saikia D.J., Croston J.H., Matsushita S., 2011, MNRAS, 417, L36

Jaffe W.J., Perola G.C., 1973, A&A, 26, 423

Jamrozy M., Konar C., Machalski J., Saikia D.J., 2008, MNRAS, 385, 1286

Jamrozy M., Machalski J., Mack K.-H., Klein U., 2005, A&A, 433, 467

Jamrozy M., Konar C., Saikia D.J., Stawarz Ł., Mack K.-H., Siemiginowska A., 2007, MNRAS, 378, 581

Jetha N.N., Hardcastle M.J., Ponman T.J., Sakelliou I., 2008, MNRAS, 391, 1052

Jones T.W., Ryu D., Engel A., 1999, ApJ, 512, 105

Kaiser C.R., 2000, A&A, 362, 447

Kaiser C.R., Schoenmakers A.P., Röttgering H.J.A., 2000, MNRAS, 315, 381

Konar C., Jamrozy M., Saikia D.J., Machalski J., 2008, MNRAS, 383, 525

Konar C., Hardcastle M.J., Croston J.H., Saikia D.J., 2009, MNRAS, 400, 480

Konar C., Saikia D.J., Jamrozy M., Machalski J., 2006, MNRAS, 372, 693

Laing, R.A., Guidetti D., Bridle A.H., Parma P., Bondi M., 2011, MNRAS, 417, 2789

Leahy J.P., Black A.R.S., Dennett-Thorpe J., Hardcastle M.J., Komissarov S., Perley R.A., Riley J.M., Scheuer P.A.G., 1997, MNRAS, 291, 20

Machalski J., Jamrozy M., Konar C., 2010, A&A, 510, 84

McNamara B.R., Nulsen P.E.J., 2007, ARA&A, 45, 117

Murgia M., 1996, Laurea Thesis, University of Bologna.

Pacholczyk A.G., 1970, Radio Astrophysics, Freeman & Co., San Francisco

Parma P., Murgia M., Morganti R., Capetti A., de Ruiter H. R., Fanti R., 1999, A&A, 344, 7

Rengelink R.B., Tang Y., de Bruyn A.G., Miley G.K., Bremer M.N., Röttgering H.J.A., Bremer M.A.R., 1997, A&AS, 124, 259

Rudnick L., Katz-Stone D., Anderson M., 1994, ApJS, 90, 955

Saikia D.J., Konar C., Kulkarni V.K., 2006, MNRAS, 366, 1391

Saikia, D.J., Jamrozy, M., 2009, BASI, 37, 63

Schoenmakers A.P., de Bruyn A.G., Röttgering H.J.A., van der Laan H., Kaiser C.R., 2000a, MNRAS, 315, 371

Schoenmakers A.P., de Bruyn A.G., Röttgering H.J.A., van der Laan H., 2000b, MNRAS, 315, 395

Spergel D.N. et al., 2003, ApJS, 148, 175

Safouris V., Subrahmanyan R., Bicknell G.V., Saripalli L., 2008, MNRAS, 385, 2117

White R.L. & Becker R.H., 1992, ApJS, 79, 331

Worrall D.M. & Birkinshaw M., 2000, ApJ, 530, 719