

Dark matter scaling relations in intermediate z haloes

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ABSTRACT

We investigate scaling relations between the dark matter (DM) halo model parameters for a sample of intermediate-redshift early-type galaxies (ETGs) resorting to a combined analysis of Einstein radii and aperture velocity dispersions. Modelling the dark halo with a Navarro–Frenk–White profile and assuming a Salpeter initial mass function (IMF) to estimate stellar masses, we find that the column density \mathcal{S} and the Newtonian acceleration within the halo characteristic radius r_s and effective radius R_{eff} are not universal quantities, but correlate with the luminosity L_V , the stellar mass M_* and the halo mass M_{200} , contrary to recent claims in the literature. We finally discuss a tight correlation among the DM mass $M_{\text{DM}}(R_{\text{eff}})$ within the effective radius R_{eff} , the stellar mass $M_*(R_{\text{eff}})$ and R_{eff} itself. The slopes of the scaling relations discussed here strongly depend, however, on the DM halo model and the IMF adopted so that these ingredients have to be better constrained in order to draw definitive conclusions on the DM scaling relations for ETGs.

Key words: galaxies: elliptical and lenticular, cD – galaxies: kinematics and dynamics – dark matter.

1 INTRODUCTION

The current cosmological paradigm, the concordance Lambda cold dark matter (Λ CDM) model, relies on two main components, namely dark energy (e.g. Carroll, Press & Turner 1992) and DM. Although in excellent agreement with all the cosmological probes (Komatsu et al. 2009; Lampteil et al. 2009; Percival et al. 2009), the Λ CDM model is nevertheless afflicted by serious problems on galactic scales. In this framework, the formation of virialized DM haloes from the initial tiny density perturbations is followed at later stages through numerical N -body simulations (Bertschinger 1998). It became apparent that the spherically averaged density profile, $\rho_{\text{DM}}(r)$, of DM haloes is independent of the halo mass (Navarro, Frenk & White 1997) and well described by a double power-law relation with $\rho_{\text{DM}} \propto r^{-3}$ in the outer regions and $\rho_{\text{DM}} \propto r^{-\alpha}$ with $\alpha > 0$ centrally. On the contrary, observations of spiral galaxies seem to definitely point towards cored models, that is, $\alpha = 0$ at the centre (de Blok 2009). Understanding whether such a discrepancy is due to some physical process not correctly modelled in simulations or due to failure of the CDM paradigm is still a hotly debated issue. As a valuable tool to address this problem, one can look for scaling relations among DM halo parameters and stellar quantities in order to better constrain the formation scenario and

the DM properties. Recently, much work has been dedicated to this issue. Using a sample of local early-type galaxies (ETGs), Tortora et al. (2009, hereafter T09) have found that DM is the main driver of the Fundamental Plane (FP) tilt (see also Cappellari et al. 2006; Bolton et al. 2007; Hyde & Bernardi 2009; Auger et al. 2010b; Graves & Faber 2010) and that the average spherical DM density is a decreasing function of stellar mass (see also Thomas et al. 2009). Based on data from rotation curves of ~ 1000 spiral galaxies, the mass models of individual dwarf and spiral galaxies and the weak lensing signal of elliptical and spirals, Donato et al. (2009, hereafter D09) and Gentile et al. (2009, hereafter G09) have found strong evidence for the constancy of the central DM column density over 12 orders of magnitude in luminosity. Napolitano, Romanowsky & Tortora (2010, hereafter NRT10) have shown that, on average, the projected density of local ETGs within effective radius is systematically higher than the same quantity for spiral and dwarf galaxies, pointing to a systematic increase with halo mass as suggested by Boyarsky et al. (2009, hereafter B09), who have extended the samples analysed above to both group and cluster scale systems.

In order to try to discriminate between these contrasting results, we present here an analysis of the DM scaling relations for a sample of ETGs at intermediate redshift ($\langle z \rangle \simeq 0.2$) using lensing and velocity dispersion data to constrain their parameters. The mass models, the data and the fitting procedure are described in Section 2. In Section 3, we describe the main results, while Section 4 is devoted to a brief review of the results and conclusions.

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2 ESTIMATING MASS QUANTITIES

As a preliminary mandatory step, we need to determine the quantities involved in the above scaling relations. To this end, one has first to choose a model for the stellar and DM components and then fit the observational data in order to infer the quantities of interest from the constrained model.

2.1 Stellar and DM profiles

Motivated by the well-known result that the surface brightness profiles of ETGs are well fitted by the Sérsic (1968) law, we describe the stellar component with the Prugniel & Simien (1997, PS hereafter) profile (see also Marquez et al. 2001 for further details). The choice of the DM halo model is quite controversial. Rotation curves of $z = 0$ spiral galaxies are better fitted by cored models (de Blok 2009), but we are here considering ETGs at intermediate z so that it is not straightforward to extend these results to our case. In order to explore the impact of the DM halo profile, we therefore adopt both a Burkert (1995) model with

$$\rho_B(r) = \frac{\rho_B r_B^3}{(r + r_B)(r^2 + r_B^2)}, \quad (1)$$

and an NFW (Navarro et al. 1997) profile with

$$\rho_{\text{NFW}}(r) = \frac{\rho_s r_s^3}{r(r + r_s)^2}. \quad (2)$$

Both the 3D and projected masses $M(r)$ and $M_{\text{proj}}(R)$ can be analytically evaluated, with $M_{\text{proj}}(R)$ given in Park & Ferguson (2003) and Bartelmann (1996) for the Burkert and NFW models, respectively.

In order to constrain the model parameters, we rely on the estimate of the projected mass $M_E = M_{\text{proj}}(R_E)$ inferred by the measurement of the Einstein radius R_E in a lens system. While lensing probes the mass projected along the line of sight, the aperture velocity dispersion σ_{ap} (Mamon & Lokas 2005) provides complementary information on the internal dynamics thus strengthening the constraints.

2.2 Data and fitting procedure

We make use of the sample of 85 lenses collected by the Sloan Lens ACS (SLACS) survey (Auger et al. 2009) and first select only ETGs with available values of both the velocity dispersion σ_{ap} (measured within an aperture of $R_{\text{ap}} = 1.5$ arcsec) and the Einstein radius R_E , thus ending up with a data set containing 59 objects. For each lens, we follow Auger et al. (2009) setting the Sérsic index $n = 4$ and the effective radius R_{eff} and total luminosity L_V to the values inferred from the V -band photometry. The SLACS collaboration has also provided an estimate of the total stellar mass (their table 4) from which we use both Chabrier (2001) and Salpeter (1955) initial mass functions (IMFs) to investigate the effect of the IMF on the scaling relations.

Before fitting the model to the data, it is worth noting that, for both the Burkert and NFW models, the halo parameters are different from one lens to another. Since we have only two observed quantities, namely $(M_E, \sigma_{\text{ap}})$ for each lens, determining (ρ_X, r_X) (with $X = B$ or s for the Burkert and NFW models, respectively) on a case-by-case basis would give us very weak constraints. We therefore bin the galaxies in 10 equally populated luminosity bins (the last one actually containing five objects) and resort to a different parametrization using quantities that it can be reasonably assumed

to be the same for all the lenses in the same bin.¹ As one of the fitting parameters, we choose the virial mass-to-light ratio (M/L) , $\Upsilon_{\text{vir}} = M_{\text{vir}}/L_V$ with M_{vir} the DM halo mass at the virial radius² R_{vir} . Should Υ_{vir} depend on L_V , our luminosity bins are quite narrow so that any change in Υ_{vir} should be so small that it can be safely neglected. We then use $\log \eta_X$ with $\eta_X = r_X/R_{\text{eff}}$ as the second parameter.

Note that, although $(\Upsilon_{\text{vir}}, \log \eta_X)$ are the same for all the galaxies in a bin, (M_{vir}, r_X) still change from one lens to another thus allowing us to estimate all the quantities of interest on a lens-by-lens basis. In order to constrain $(\Upsilon_{\text{vir}}, \log \eta_X)$, we maximize a suitable likelihood function, composed of two terms, the first (second) one referring to the lensing (dynamics) constraints (see Cardone et al. 2009 for further details). In order to efficiently explore the parameter space, we use a Markov chain Monte Carlo (MCMC) algorithm running chains with 100 000 points reduced to more than 3000 after burn-in cut and thinning. For a given galaxy, we compute the quantities of interest for each point in the chains and then use Bayesian statistics to infer median values and 68 per cent confidence intervals. We follow D'Agostini (2004) to correct for the asymmetric errors in our estimates.

3 RESULTS

Before investigating scaling relations, it is worth checking whether our PS + DM model works in fitting the lens data.

3.1 The fiducial stellar + DM profile

We have considered four PS + DM models by combining the two halo profiles (Burkert or NFW) with the two IMFs (Chabrier or Salpeter) adopted to set the total stellar mass. We find that all the four models fit remarkably well the lensing and dynamical data with rms deviations $\sim 1\sigma$ and the lens observed values deviating no more than 2σ from the model ones for most of the bins. On the one hand, such a result ensures us that the estimates of the different quantities entering the scaling relations we are interested in are based on empirically motivated models. On the other hand, this same result tells us that the data we are using are unable to discriminate among different choices. This is an outcome of the limited radial range the data probe. Indeed, given typical values of R_E/R_{eff} and $R_{\text{ap}}/R_{\text{eff}}$, both M_E and σ_{ap} mainly probe the inner regions of the lenses so that they can provide only weak constraints on the behaviour of the mass profile in the outer DM halo dominated regions.

Although statistically equivalent (as can be quantitatively judged on the basis of the close values of \mathcal{L}_{max} in Tables 1 and 2), the four models may be ranked by examining the constraints on their parameters. Indeed, the values in Table 1 show that the PS + Burkert models fit the data with quite low values of Υ_{vir} thus leading to unexpectedly small virial masses. Roughly averaging the maximum likelihood Υ_{vir} of the different bins, we get $\langle \Upsilon_{\text{vir}} \rangle = 22 \pm 14$ (12 ± 9) $M/L_{V,\odot}$ using the Chabrier (Salpeter) IMF. On the contrary, when adopting the NFW profile, we find $\langle \Upsilon_{\text{vir}} \rangle = 590 \pm 210$ (360 ± 150) $M/L_{V,\odot}$ for a Chabrier (Salpeter) IMF. Our results

¹Note that this procedure allows us to fit for these two alternative parameters on a bin-by-bin basis thus having $2\mathcal{N}_{\text{bin}} - 2$ degrees of freedom with \mathcal{N}_{bin} the number of lenses in the bin.

²We define R_{vir} as the radius where the mean density $M_{\text{vir}}/(4/3)\pi R_{\text{vir}}^3$ equals $\Delta_c(z)\bar{\rho}_M(z)$ with $\Delta_c(z)$ as in Bryan & Norman (1998) and $\bar{\rho}_M$ the mean matter density at z .

Table 1. Constraints on halo parameters for the Burkert models with Chabrier and Salpeter IMF. The median luminosity from the galaxies in each bin is reported in column 1, while the adjacent columns give the maximum likelihood parameter ($\Upsilon_{\text{vir}}, \log \eta_s$)_{ML}, the value of $-2 \ln \mathcal{L}$ at the maximum, median value and 68 per cent confidence interval for the DM parameters ($\Upsilon_{\text{vir}}, \log \eta_s$) for each model.

Bin	Chabrier IMF				Salpeter IMF			
	($\Upsilon_{\text{vir}}, \log \eta_s$) _{ML}	$-2 \ln \mathcal{L}_{\text{max}}$	(Υ_{vir}) _{-1σ} ^{+1σ}	($\log \eta_s$) _{-1σ} ^{+1σ}	($\Upsilon_{\text{vir}}, \log \eta_s$) _{ML}	$-2 \ln \mathcal{L}_{\text{max}}$	(Υ_{vir}) _{-1σ} ^{+1σ}	($\log \eta_s$) _{-1σ} ^{+1σ}
10.72	(15, -0.61)	353.179	26 ⁺²² ₋₁₂	-0.38 ^{+0.21} _{-0.24}	(15, -0.80)	351.645	19 ⁺³² ₋₁₁	-0.24 ^{+0.19} _{-0.32}
10.82	(53, -0.18)	362.797	46 ⁺⁶⁰ ₋₃₂	-0.22 ^{+0.22} _{-0.22}	(35, -0.15)	364.679	41 ⁺³⁰ ₋₂₃	-0.09 ^{+0.15} _{-0.25}
10.85	(9, -0.81)	367.780	14 ⁺¹¹ ₋₆	-0.62 ^{+0.22} _{-0.21}	(15, -0.80)	364.975	14 ⁺⁴² ₋₈	-0.36 ^{+0.38} _{-0.31}
10.95	(32, -0.37)	364.059	43 ⁺⁴⁸ ₋₂₁	-0.29 ^{+0.20} _{-0.21}	(17, -0.41)	364.379	29 ⁺³³ ₋₁₆	-0.25 ^{+0.21} _{-0.25}
11.00	(27, -0.31)	366.641	36 ⁺¹⁹ ₋₁₆	-0.21 ^{+0.17} _{-0.20}	(15, -0.26)	366.721	37 ⁺⁵⁴ ₋₂₃	0.02 ^{+0.23} _{-0.32}
11.07	(10, -0.80)	375.608	13 ⁺⁸ ₋₅	-0.26 ^{+0.18} _{-0.23}	(10, -0.87)	372.049	11 ⁺²² ₋₅	-0.54 ^{+0.28} _{-0.21}
11.12	(24, -0.41)	370.943	38 ⁺³⁴ ₋₁₈	-0.26 ^{+0.18} _{-0.23}	(14, -0.40)	370.400	40 ⁺⁶⁰ ₋₂₄	-0.07 ^{+0.23} _{-0.29}
11.15	(20, -0.43)	377.986	39 ⁺²⁷ ₋₂₀	-0.23 ^{+0.19} _{-0.21}	(8, -0.58)	377.817	18 ⁺³³ ₋₁₀	-0.26 ^{+0.28} _{-0.28}
11.27	(16, -0.61)	388.842	23 ⁺²³ ₋₁₆	-0.48 ^{+0.25} _{-0.21}	(10, -0.60)	384.703	26 ⁺³⁴ ₋₁₅	-0.65 ^{+0.35} _{-0.29}
11.43	(8, -1.05)	319.962	13 ⁺¹⁵ ₋₅	-0.78 ^{+0.30} _{-0.25}	(9, -1.10)	319.953	12 ⁺²³ ₋₅	-0.65 ^{+0.35} _{-0.29}

Table 2. Same as Table 1, but for NFW models.

Bin	Chabrier IMF				Salpeter IMF			
	($\Upsilon_{\text{vir}}, \log \eta_s$) _{ML}	$-2 \ln \mathcal{L}_{\text{max}}$	(Υ_{vir}) _{-1σ} ^{+1σ}	($\log \eta_s$) _{-1σ} ^{+1σ}	($\Upsilon_{\text{vir}}, \log \eta_s$) _{ML}	$-2 \ln \mathcal{L}_{\text{max}}$	(Υ_{vir}) _{-1σ} ^{+1σ}	($\log \eta_s$) _{-1σ} ^{+1σ}
10.72	(1697, 1.46)	352.771	1408 ⁺⁶⁴¹ ₋₄₃₄	1.39 ^{+0.14} _{-0.14}	(2074, 1.92)	350.930	713 ⁺⁶²⁵ ₋₃₄₀	1.48 ^{+0.22} _{-0.23}
10.82	(996, 1.24)	364.487	968 ⁺⁵⁵⁵ ₋₂₅₄	1.23 ^{+0.18} _{-0.13}	(590, 1.30)	365.630	547 ⁺⁵²⁶ ₋₂₂₈	1.28 ^{+0.31} _{-0.22}
10.85	(544, 1.01)	366.455	544 ⁺¹²² ₋₈₆	1.03 ^{+0.09} _{-0.08}	(326, 1.15)	363.877	369 ⁺³⁸⁰ ₋₁₂₁	1.23 ^{+0.25} _{-0.21}
10.95	(544, 1.00)	365.776	551 ⁺¹⁰⁶ ₋₈₉	1.00 ^{+0.09} _{-0.07}	(322, 1.05)	365.268	320 ⁺¹²⁹ ₋₆₇	1.07 ^{+0.17} _{-0.12}
11.00	(882, 1.34)	366.180	795 ⁺⁴⁸⁶ ₋₂₅₀	1.31 ^{+0.20} _{-0.16}	(534, 1.52)	365.707	491 ⁺⁷⁷⁹ ₋₂₄₈	1.52 ^{+0.43} _{-0.31}
11.07	(428, 0.93)	377.148	432 ⁺⁷⁵ ₋₆₁	0.94 ^{+0.07} _{-0.06}	(233, 0.96)	373.223	235 ⁺⁶⁵ ₋₅₀	0.99 ^{+0.14} _{-0.11}
11.12	(511, 1.07)	373.473	525 ⁺¹³⁶ ₋₉₃	1.09 ^{+0.19} _{-0.10}	(267, 1.09)	371.347	262 ⁺⁹¹ ₋₆₆	1.10 ^{+0.15} _{-0.14}
11.15	(383, 0.99)	373.643	391 ⁺⁶⁹ ₋₅₉	1.00 ^{+0.07} _{-0.06}	(183, 0.98)	374.078	203 ⁺⁶⁵ ₋₅₀	1.02 ^{+0.18} _{-0.11}
11.27	(503, 1.05)	394.857	513 ⁺¹²⁰ ₋₁₀₀	1.05 ^{+0.10} _{-0.11}	(274, 1.07)	388.876	286 ⁺¹³⁶ ₋₇₃	1.11 ^{+0.18} _{-0.15}
11.43	(529, 0.96)	322.021	532 ⁺⁸⁶ ₋₇₆	0.97 ^{+0.06} _{-0.07}	(529, 0.96)	322.015	532 ⁺⁸⁶ ₋₇₇	0.97 ^{+0.07} _{-0.07}

from NFW + Salpeter are qualitatively consistent with previous estimates of $\Upsilon_{\text{vir}} \sim 200 M/L_{V,\odot}$ obtained relying on galaxy–galaxy lensing (Guzik & Seljak 2002; Hoekstra et al. 2005) or combining strong and weak lensing and central dynamics (Gavazzi et al. 2007). Therefore, we will consider the NFW + Salpeter as our fiducial model,³ retaining the other cases just for investigating the dependence of the scaling relations on the halo profiles and IMF adopted. This assumption is also confirmed by recent findings which point to a Salpeter IMF when an uncontracted NFW profile is assumed (Auger et al. 2010a; NRT10; Treu et al. 2010).

The constraints on ($\Upsilon_{\text{vir}}, \log \eta_s$) for the four cases considered are summarized in Tables 1 and 2 for Burkert and NFW models, respectively. The marginalized constraints on Υ_{vir} show that the NFW + Salpeter model provides reasonable values for the virial M/L in all bins (but the first) thus motivating our choice as a fiducial case. The 68 per cent CL for Υ_{vir} are, however, quite large and asymmetrically extended towards very large values. This can

be qualitatively explained noting that the larger is Υ_{vir} , the larger is M_{vir} and hence R_{vir} . In such a case, R_E/R_{vir} and $R_{\text{ap}}/R_{\text{vir}}$ become quite small such that the data are less and less able to constrain the outer regions. As a consequence, the marginalized likelihood function $\mathcal{L}(\Upsilon_{\text{vir}})$ has a long flat tail for large Υ_{vir} thus giving rise to the reported asymmetric errors. A similar effect also explains why the errors on $\log \eta_s$ are asymmetric and still large. The wide confidence intervals prevent us from quantifying how Υ_{vir} and $\log \eta_s$ change from one bin to another. Excluding the first bin, characterized by large uncertainties, we find that both ($\Upsilon_{\text{vir}}, \log \eta_s$) do not show any significant trend with L_V .

3.2 DM correlations

We now investigate the relation between the column density $\mathcal{S}(R) = M_{\text{proj}}(<R)/\pi R^2$, luminosity, stellar and halo mass⁴ for the lenses in our sample. The main results of this analysis are shown in Fig. 1. Despite the small mass range probed and the large errors, we find that $\mathcal{S}(r_s)$ is positively correlated with L_V, M_* and M_{200} , confirming the results in B09. The slope of the correlations may be

³ In order to put this choice on firmer statistical grounds, we could have redefined our likelihood by adding a prior on Υ_{vir} rather than the uninformative flat one we have used here. However, since previous estimates of Υ_{vir} are affected by large errors and based on model assumptions, we have preferred to avoid the risk of a fit driven by the prior examining a posteriori the resulting values.

⁴ Note that, in order to be consistent with B09, we use M_{200} , that is, the mass within the radius R_{200} , where the mean density is 200 times the cosmological mean matter density.

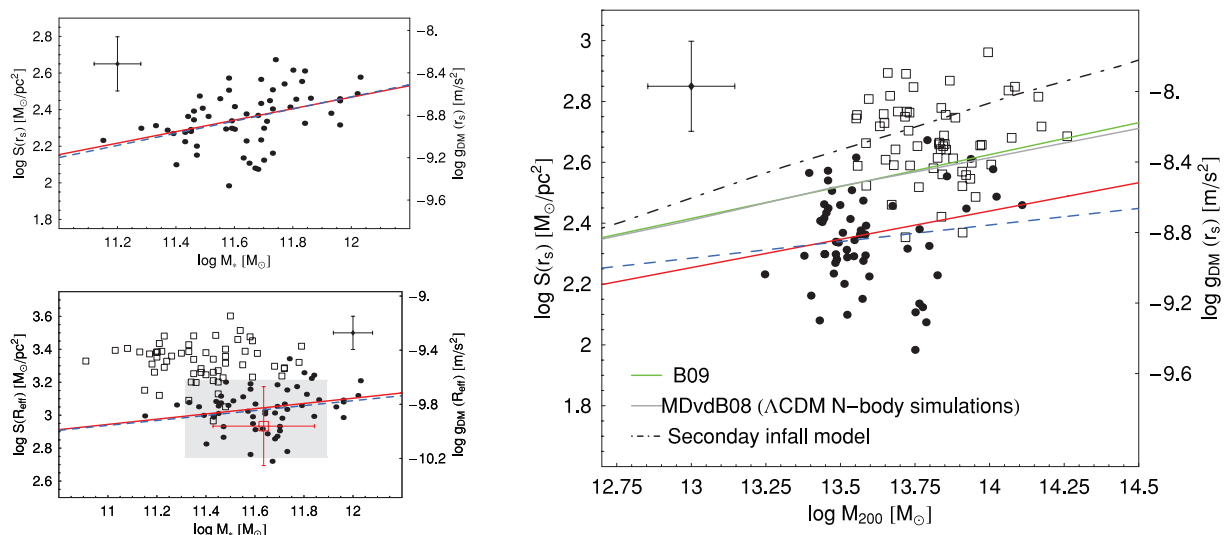


Figure 1. $S(r_s)$ and $S(R_{\text{eff}})$ for the NFW + Salpeter model as a function of stellar mass (left-hand panels) and $S(r_s)$ versus halo mass M_{200} (right-hand panel), shown as black points. We superimpose the best-fitting linear relation using the D’Agostini (2005) (red solid) and direct fit (blue dashed) methods. On the right-hand axes is shown the equivalent acceleration scale g_{DM} , defined as in the text and derived by the fit of \mathcal{S} – g_{DM} relation. The black error bars set the median data uncertainties. The results when a Chabrier IMF is used are shown as open boxes. In the left-bottom panel, the median from NRT10 and Tortora et al. (2010) are shown as red symbols and grey region, respectively (in both the cases the 1σ scatter of the distribution is plotted). In the right-hand panel, we add the B09 best-fitting linear relation, the secondary-infall model (Boyarisky et al. 2010) prediction and the results from the Λ CDM N -body simulation of Macciò et al. (2008).

estimated using the D’Agostini (2005) fitting method, which takes into account the errors on both variables and the intrinsic scatter σ_{int} . Concentrating on the trend with stellar mass, for the maximum likelihood parameters,⁵ we get

$$\log S(r_s) = 0.29 \log \left(\frac{M_*}{10^{11} M_\odot} \right) + 2.17$$

with an intrinsic scatter⁶ $\sigma_{\text{int}} = 0.01$. Although the results from a direct fit⁷ are only slightly different, for completeness we will plot the best-fitting relations from both the methods (see Fig. 1). Similar results are found when the column density is fitted as a function of L_V .

If we fit the same relation, but replace $S(r_s)$ with $S(R_{\text{eff}})$, the trends are shallower, and the zero-points change too. As shown in Fig. 1, an error-weighted average over the galaxies in the sample gives $\langle \log S(R_{\text{eff}}) \rangle \sim 3.1$ in agreement, within the scatter, with the median $\log S(R_{\text{eff}}) \simeq 2.9$ from the local ETG sample of T09 and NRT10 and the results in Tortora et al. (2010) where a different analysis, based on an isothermal profile, is used on the same lens sample. All these results are in agreement with a scenario where ETGs surface from the merging of late-type systems so that, at a fixed radius, their DM content is larger than the one for spirals

⁵ We will refer always to the maximum likelihood values only, but the reader must be aware that, in a Bayesian framework, what is most important are the marginalized values and their confidence ranges given in Table 3. Since the median value of the slope and the scatter are close to the maximum likelihood ones, this choice has no impact on the discussion.

⁶ The intrinsic scatter accounts for the deviations of the single galaxies from the underlying model leading to the fitted relation.

⁷ In the direct fit, we minimize the usual χ^2 assuming that the errors on the x variable are negligible and no intrinsic scatter is present. These simplifying assumptions do not hold for most of the scaling relations we have considered so that we resort to the D’Agostini (2005) method as a more reliable procedure. However, if σ_{int} is small and the error on x lower than that on y , the two methods converge to the same maximum likelihood values.

Table 3. Marginalized constraints on the scaling relation parameters for the correlations involving the column density \mathcal{S} assuming the fiducial NFW + Salpeter model. Columns are as follows: (1) correlation ID; (2–4) median value and 68 confidence ranges for $(\alpha, \beta, \sigma_{\text{int}})$, where the linear relation $\log y = \alpha \log x + \beta$ is fitted and σ_{int} is the intrinsic scatter. Note that, due to the fitting method, for each value of $(\alpha, \sigma_{\text{int}})$, β is set by the condition that the likelihood is maximized, in other words we analytically marginalize over β when determining the maximum likelihood parameters. We warn the reader that, in the fit, we use the luminosity L_V in units of $10^{11} L_\odot$ and the stellar (halo) mass in units of 10^{11} (10^{12}) M_\odot to reduce error covariance.

ID	$(\alpha)_{-1\sigma}^{+1\sigma}$	$(\beta)_{-1\sigma}^{+1\sigma}$	$(\sigma_{\text{int}})_{-1\sigma}^{+1\sigma}$
$\log S(r_s) - \log L_V$	$0.28_{-0.13}^{+0.13}$	$2.34_{-0.01}^{+0.01}$	$0.037_{-0.027}^{+0.043}$
$\log S(r_s) - \log M_*$	$0.29_{-0.15}^{+0.15}$	$2.16_{-0.09}^{+0.10}$	$0.037_{-0.027}^{+0.043}$
$\log S(r_s) - \log M_{200}$	$0.14_{-0.15}^{+0.15}$	$2.11_{-0.22}^{+0.24}$	$0.040_{-0.029}^{+0.044}$
$\log S(R_{\text{eff}}) - \log M_*$	$0.14_{-0.12}^{+0.12}$	$2.94_{-0.08}^{+0.07}$	$0.057_{-0.032}^{+0.039}$

and dwarves [$\log S(R_{\text{eff}}) \sim 2$ – 2.5 , see NRT10 for further details]. Should we use a Chabrier IMF (thus lowering the stellar masses by a factor of ~ 1.8) we get a larger DM content and $S(R_{\text{eff}})$ which is a constant function of M_* (see open boxes in Fig. 1).

A weaker positive correlation is found when plotting $S(r_s)$ versus M_{200} , the maximum likelihood fit being

$$\log S(r_s) = 0.16 \log \left(\frac{M_{200}}{10^{12} M_\odot} \right) + 2.11$$

with $\sigma_{\text{int}} = 0.010$. Our best-fitting relation is shallower than the B09 one, although the slope is consistent with their one (0.21) within the large error bars. Note that we have here explicitly taken into account the correlated errors on $S(r_s)$ and M_{200} , while we do not know whether the fitting method adopted by B09 does the same. We therefore cannot exclude that the difference in slope is only an outcome of the use of different algorithms on noisy data. Actually, our $S(r_s)$ values are on average smaller than those in B09

(over the same mass range). Our results are also systematically smaller than the estimates from the Λ CDM N -body simulations of isolated haloes from Macciò, Dutton & van den Bosch (2008) and the predictions from the secondary infall model⁸ (B09; Del Popolo 2009). Although a wider mass range has to be probed to infer any definitive answer, we argue that the larger values in the literature are a consequence of neglecting the stellar component. The agreement with B09 would improve if a Chabrier IMF was used (see open boxes in the right panel of Fig. 1) and, as discussed above, the trend would flatten out.

When discussing the results for our reference model, the observed correlations argue against the universality of the column density proposed in D09.⁹ It is worth investigating why we and D09 reach completely opposite conclusions. As a first issue, we note that D09 describe the DM halo adopting the Burkert model. Should we use this model to infer $\mathcal{S}(1.6r_B)$, a Chabrier IMF (which does not strongly differ from the IMFs used in D09), and plot as a function of luminosity to be uniform with them, the best-fitting relation would have been

$$\log \mathcal{S}(1.6r_B) = 0.02 \log \left(\frac{L_V}{10^{11} L_\odot} \right) + 2.65$$

with $\sigma_{\text{int}} = 0.15$. In such a case, within a very good approximation, we can assume $\mathcal{S}(1.6r_B)$ is indeed constant with L_V . An error-weighted average of the sample values gives $\langle \log \mathcal{S}(1.6r_B) \rangle = 3.07 \pm 0.02 \pm 0.20$, where here and in following similar estimates, the first error is the statistical uncertainty and the second one the rms around the mean. Taken at face value, this estimate is significantly larger than the D09 one (similarly to what was found when discussing the comparison among the $\mathcal{S}(R_{\text{eff}})$ values for ETGs and spirals), even if it is marginally within $2\sigma_{\text{rms}}$. However, as discussed above, the slope of the $\mathcal{S}_B(1.6r_B) - L_V$ relation is strongly model dependent. Indeed, changing to a Salpeter IMF, we find $\mathcal{S}_B(1.6r_B) \propto L_V^{0.20}$ thus arguing in favour of non-universality. Actually, fitting the Burkert model to the data gives viral masses that are likely too small. As such, the scaling relations obtained using this model are likely to be biased and one should rely more on the results for the Salpeter + NFW model, thus concluding that the column density \mathcal{S} indeed depends on the mass for intermediate-redshift ETGs. Investigating the reasons why NFW models are preferred over Burkert ones for ETGs, while the opposite is true for spirals can provide important constraints on galaxy formation scenarios, but it is outside our scope here.

Motivated by the D09 findings on the universality of $\rho_B r_B$ and noting that the DM Newtonian acceleration, $g_{\text{DM}}(r) = GM_{\text{DM}}(r)/r^2$, is proportional to $\rho_B r_B$ for a Burkert model, G09 have recast the D09 results in terms of a constancy of $g_{\text{DM}}(r_B)$. Then they extended this result showing that also the stellar acceleration, $g_*(r) = GM_*(r)/r^2$ at r_B is a universal quantity. Using their same Burkert model and a Chabrier IMF, we indeed find that $g_{\text{DM}}(r_B)$ is roughly constant with L_V , with an error-weighted value $\langle \log g_{\text{DM}}(r_B) \rangle = -8.89 \pm 0.02 \pm 0.20$, in satisfactory agreement with the G09 result [$\log g_{\text{DM}}(r_B) \simeq$

-8.5]. Having said that, we find that $g_*(r_B)$ is not constant, the best-fitting relation predicting $g_*(r_B) \propto L_V^{-0.21}$. We stress, however, that the situation is completely reversed if we use a Salpeter IMF giving $g_{\text{DM}}(r_B) \propto L_V^{0.20}$ and $g_*(r_B) \propto L_V^{-0.01}$, so that drawing a definitive answer on which quantity is universal is an ambiguous task. Whatever is the correct IMF, we can nevertheless safely conclude that the DM and stellar Newtonian accelerations cannot both be universal quantities, in contrast with the claim in G09. Moreover, should the IMF be universal and intermediate between Chabrier and Salpeter, one could argue that neither $g_{\text{DM}}(r_B)$ and $g_*(r_B)$ are universal quantities.

As already discussed, our fiducial model is the NFW + Salpeter so that it is preferable to use this model when investigating whether the Newtonian accelerations are constant or not. Moreover, rather than using the halo characteristic radius (which refers to a different mass content depending on the model), we will discuss the results at R_{eff} thus referring to the better constrained inner regions. For the fiducial model, we find $\log g_{\text{DM}}(R_{\text{eff}}) \propto 0.26 \log L_V$ thus arguing against the universality of this quantity. However, as can be seen in Table 4, the slope of the $g_{\text{DM}}(R_{\text{eff}}) - L_V$ correlation is strongly model dependent with values indicating either an increasing behaviour (for NFW models), a flat one (for Burkert + Salpeter) or a decreasing one (for Burkert + Chabrier). These trends are expected from the analysis of column density made above. On the contrary, the scaling of $g_*(R_{\text{eff}})$ does not depend on the halo model because its value is estimated from stellar quantities only, while the effect of the IMF is simply a systematic rescaling due to the higher stellar masses for the Salpeter case. Although the slope in Table 4 points towards a significant decrease with the luminosity, it is worth stressing that assuming a constant value ($\log g_*(R_{\text{eff}}) = -9.403 \pm 0.01 \pm 0.20$) (for a Salpeter IMF) provides a similarly good fit so that we cannot draw a definitive answer.

Any correlation between a DM quantity and a stellar one may be the outcome of a hidden interaction between the two galactic components. In particular, for the Newtonian accelerations, because of their definition, it is straightforward to show that $g_{\text{DM}}(R_{\text{eff}}) = GM_{\text{DM}}(R_{\text{eff}})/R_{\text{eff}}^2 = M_{\text{DM}}(R_{\text{eff}})/M_{\text{eff}} g_*(R_{\text{eff}})$ with $M_{\text{eff}} = M_*(R_{\text{eff}})$ so that one can look for a correlation between these two quantities. For the NFW + Salpeter model, we indeed get

$$\log g_{\text{DM}}(R_{\text{eff}}) = 0.21 \log g_*(R_{\text{eff}}) - 7.89.$$

Table 4. Marginalized constraints on the scaling relation parameters for the correlations $\log g_{\text{DM}}(R_{\text{eff}}) - \log L_V$ (upper part) and $\log g_*(R_{\text{eff}}) - \log L_V$ (lower part) with the accelerations in m s^{-2} and the luminosity in units of $10^{11} L_\odot$. Columns are as follows: (1) model ID, (2–4) median value and 68 per cent confidence ranges for $(\alpha, \beta, \sigma_{\text{int}})$. The model IDs are BSC for Burkert + Chabrier, BSS for Burkert + Salpeter, NSC for NFW + Chabrier, NSS for NFW + Salpeter. Note that since $g_*(R_{\text{eff}})$ refers to stellar quantities only, the correlations involving this quantity are independent of the halo model.

Model ID	$(\alpha)_{-1\sigma}^{+1\sigma}$	$(\beta)_{-1\sigma}^{+1\sigma}$	$(\sigma_{\text{int}})_{-1\sigma}^{+1\sigma}$
BSC	$-0.155_{-0.145}^{+0.145}$	$-0.090_{-0.003}^{+0.003}$	$0.175_{-0.019}^{+0.023}$
BSS	$-0.004_{-0.142}^{+0.144}$	$-9.322_{-0.002}^{+0.004}$	$0.156_{-0.029}^{+0.033}$
NSC	$0.088_{-0.111}^{+0.111}$	$-9.526_{-0.004}^{+0.004}$	$0.066_{-0.035}^{+0.043}$
NSS	$0.245_{-0.125}^{+0.126}$	$-9.841_{-0.006}^{+0.007}$	$0.040_{-0.028}^{+0.043}$
NSC	$-0.461_{-0.168}^{+0.170}$	$-9.649_{-0.005}^{+0.004}$	$0.178_{-0.022}^{+0.028}$
NSS	$-0.461_{-0.168}^{+0.170}$	$-9.403_{-0.005}^{+0.004}$	$0.178_{-0.022}^{+0.028}$

⁸ Note that the predictions for the secondary infall model of Del Popolo are actually smaller than the Boyarsky et al. (2010) model ones (Del Popolo et al., in preparation).

⁹ Actually, what D09 refers to as universal quantity is the product $\rho_B r_B$. It is, however, possible to show that $\mathcal{S}_{\text{NFW}}(r_s) \simeq \mathcal{S}_B(1.6r_B)$ (B09). It is then just a matter of algebra to get $\mathcal{S}_B(1.6r_B) \simeq 1.89 \rho_B r_B$ so that the constraint $\log(\rho_B r_B) = 2.17 \pm 0.20$ in D09 translates into $\log \mathcal{S}_B(1.6r_B) = 2.44 \pm 0.47$ which we use as a comparison value.

With $\sigma_{\text{int}} = 0.018$, the marginalized constraints are

$$\alpha = 0.20_{-0.13}^{+0.13}, \quad \beta = -7.90_{-1.28}^{+1.13}, \quad \sigma_{\text{int}} = 0.042_{-0.030}^{+0.043}.$$

Actually, we can recast the above relation in a different way. From the definitions of $g_{\text{DM}}(R_{\text{eff}})$ and $g_{\star}(R_{\text{eff}})$ and the assumption $\log g_{\text{DM}}(R_{\text{eff}}) \propto \alpha \log g_{\star}(R_{\text{eff}})$, one easily gets

$$\log M_{\text{DM}}(R_{\text{eff}}) \propto 2(1 - \alpha) \log R_{\text{eff}} + \alpha \log M_{\text{eff}}$$

so that we fit a log-linear relation

$$\log M_{\text{DM}}(R_{\text{eff}}) = \alpha_{\text{M}} \log R_{\text{eff}} + \beta_{\text{M}} \log M_{\text{eff}} + \gamma_{\text{M}}.$$

For the best-fitting relation, we find (1.46, 0.60, 2.79) with $\sigma_{\text{int}} = 0.006$, while the marginalized constraints (median and 68 per cent CL) read

$$\alpha_{\text{M}} = 1.47_{-0.26}^{+0.25}, \quad \beta_{\text{M}} = 0.61_{-0.21}^{+0.22},$$

$$\gamma_{\text{M}} = 2.54_{-0.26}^{+0.41}, \quad \sigma_{\text{int}} = 0.029_{-0.020}^{+0.033}.$$

For $\alpha = 0.21$, we expect $\alpha_{\text{M}} \simeq 1.58$ in agreement with our estimate. On the contrary, β_{M} is significantly larger than α possibly indicating that the ratio $g_{\text{DM}}(R_{\text{eff}})/g_{\star}(R_{\text{eff}})$ depends on the stellar mass more than expected. However, because of the correlation between β_{M} and σ_{int} induced by the fit, a wrong estimate of the intrinsic scatter may induce a bias in the best fit β_{M} . Since the error bars are quite large, determining σ_{int} is a difficult task as can also be understood noting that the best-fitting σ_{int} is formally outside the marginalized 68 per cent confidence range (but within the 95 per cent one). Nevertheless, the small value of the rms of the residuals ($\sigma_{\text{rms}} = 0.12$) is strong evidence in favour of a very tight correlation. We also stress that this scaling relation (although with different coefficients) still holds if we change the IMF or the halo model. Interestingly, this correlation is pretty similar to the luminosity and mass FP discussed in Bolton et al. (2007) and Hyde & Bernardi (2009), with the total M/L found to depend less on the stellar mass density than on the effective radius.

An alternative way to look for the correlation between the stellar and DM mass at R_{eff} may be provided by the DM mass fraction, $f_{\text{DM}}(R_{\text{eff}}) = M_{\text{DM}}(R_{\text{eff}})/[M_{\star}(R_{\text{eff}}) + M_{\text{DM}}(R_{\text{eff}})]$. Since $M_{\text{DM}}(R_{\text{eff}})$ and $M_{\star}(R_{\text{eff}})$ are correlated, we expect to find a similar correlation between $f_{\text{DM}}(R_{\text{eff}})$ and mass proxies, such as M_{\star} and L_V (Cappellari et al. 2006; Bolton et al. 2007; Hyde & Bernardi 2009, T09; Auger et al. 2010b). We indeed find

$$\log f_{\text{DM}}(R_{\text{eff}}) = 0.49 \log \left(\frac{M_{\star}}{10^{11} L_{\odot}} \right) - 0.86,$$

$$\log f_{\text{DM}}(R_{\text{eff}}) = 0.51 \log \left(\frac{L_V}{10^{11} L_{\odot}} \right) - 0.56,$$

with $\sigma_{\text{int}} = 0.03$ (0.02) and $\sigma_{\text{rms}} = 0.13$ (0.12) for the first (second) case. The marginalized constraints for the $f_{\text{DM}}-M_{\star}$ relation are

$$\alpha = 0.48_{-0.11}^{+0.11}, \quad \beta = -0.86_{-0.08}^{+0.08}, \quad \sigma_{\text{int}} = 0.049_{-0.032}^{+0.041},$$

while for the $f_{\text{DM}}-L_V$ we find

$$\alpha = 0.51_{-0.09}^{+0.09}, \quad \beta = -0.57_{-0.01}^{+0.01}, \quad \sigma_{\text{int}} = 0.028_{-0.021}^{+0.038},$$

The large error bars on the individual points likely make the estimate of σ_{int} biased, but we nevertheless find clear evidence for a DM content increasing with both M_{\star} and L_V . Both these correlations are in very good agreement with what is found in T09 for local ETGs for high-luminosity systems ($\log L_B \geq 10.4$), notwithstanding the different models adopted (NFW halo and Sersic stellar profile here versus a full isothermal model in T09). However, the slopes we find are strongly dependent on the model assumptions. Should we use the same NFW model, but the Chabrier IMF, we get

$f_{\text{DM}}(R_{\text{eff}}) \propto L_V^{0.20}$ and $f_{\text{DM}}(R_{\text{eff}}) \propto M_{\star}^{0.16}$, which are fully consistent with the results we have obtained in Cardone et al. (2009), where a general galaxy model has been fitted to a subsample of SLACS lenses.¹⁰

4 CONCLUSIONS

Much attention has recently been dedicated to investigating whether some correlations can be found among DM quantities and the stellar ones with contrasting results pointing towards a universal DM column density \mathcal{S} (D09; G09) or its variation with halo mass M_{200} (B09) and luminosity M_{\star} (T09; NRT10; Auger et al. 2010a). Here we have addressed this controversy using a sample of intermediate-redshift ETGs using the available data on both the projected mass within the Einstein radius and the aperture velocity dispersion to fit four different stellar + DM halo models. Motivated by recent findings in the literature (Treu et al. 2010; NRT10) and the analysis of the virial M/L , we have finally chosen a NFW DM halo and a Salpeter IMF which we have then used as a reference case for investigating the scaling relations of interest.

Contrary to D09, we find that the column density \mathcal{S} , evaluated at both the halo characteristic radius r_s , or the stellar effective radius R_{eff} is not a universal quantity, but rather correlates with the luminosity L_V and the stellar and halo masses M_{\star} and M_{200} . Although the slopes of these correlations depend on the halo model and IMF, assuming our reference model, the $\mathcal{S}(r_s)-M_{200}$ relation we find agrees with the B09 one, with a rather similar slope (0.19 versus 0.21), but a smaller zero-point. As a consequence, our $\mathcal{S}(r_s)$ values are smaller than the B09 ones, but also smaller than those predicted on the basis of a secondary infall model (B09; Del Popolo 2009) and Λ CDM N -body simulations (Macciò et al. 2008). We argue that this discrepancy is expected considering that these studies do not add a stellar component to the galaxy model, while here we take this explicitly into account thus decreasing the DM content. We have also found that the ensemble average column density in the central regions is systematically larger than the one in spiral galaxies in agreement with, e.g. NRT10. This is consistent with mass accretion in more massive haloes due to merging of late-type systems. As an interesting new result, we have shown that a very tight log-linear relation among $M_{\text{DM}}(R_{\text{eff}})$, R_{eff} and M_{eff} can be found leading to a DM mass fraction which positively correlates with both the stellar luminosity and mass (Cappellari et al. 2006; Bolton et al. 2007; Hyde & Bernardi 2009; T09).

The limited mass and luminosity range probed and the large errors on the different quantities involved prevent us from drawing a definitive answer on the slope and normalization of the above scaling relations. Moreover, a larger data set should also allow us to make a more detailed investigation of the impact of halo profiles and IMF assumptions. Should these further tests confirm our results, DM scaling relations can provide a valuable tool in understanding the physical processes which drive galaxy formation and evolution.

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¹⁰ Note that there is a typo in Cardone et al. (2009), the best-fitting relation being $f_{\text{DM}}(R_{\text{eff}}) \propto M_{\star}^{0.13}$, while the normalization refers to a Chabrier IMF.

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