

NIR Tully–Fisher in the Zone of Avoidance – III. Deep NIR catalogue of the HIZOA galaxies

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Accepted 2016 July 27. Received 2016 July 27; in original form 2016 June 27

ABSTRACT

We present a deep near-infrared (NIR; J , H , and K_s bands) photometric catalogue of sources from the Parkes H I Zone of Avoidance (HIZOA) survey, which forms the basis for an investigation of the matter distribution in the Zone of Avoidance. Observations were conducted between 2006 and 2013 using the Infrared Survey Facility (IRSF), a 1.4-m telescope situated at the South African Astronomical Observatory site in Sutherland. The images cover all 1108 HIZOA detections and yield 915 galaxies. An additional 105 bright 2MASS galaxies in the southern ZOA were imaged with the IRSF, resulting in 129 galaxies. The average K_s -band seeing and sky background for the survey are 1.38 arcsec and 20.1 mag, respectively. The detection rate as a function of stellar density and dust extinction is found to depend mainly on the H I mass of the H I detected galaxies, which in principal correlates with the NIR brightness of the spiral galaxies. The measured isophotal magnitudes are of sufficient accuracy (errors ~ 0.02 mag) to be used in a Tully–Fisher analysis. In the final NIR catalogue, 285 galaxies have both IRSF and 2MASS photometry (180 HIZOA plus 105 bright 2MASX galaxies). The K_s -band isophotal magnitudes presented in this paper agree, within the uncertainties, with those reported in the 2MASX catalogue. Another 30 galaxies, from the HIZOA northern extension, are also covered by UKIDSS Galactic Plane Survey (GPS) images, which are one magnitude deeper than our IRSF images. A modified version of our photometry pipeline was used to derive the photometric parameters of these UKIDSS galaxies. Good agreement was found between the respective K_s -band isophotal magnitudes. These comparisons confirm the robustness of the isophotal parameters and demonstrate that the IRSF images do not suffer from foreground contamination, after star removal, nor underestimate the isophotal fluxes of ZoA galaxies.

Key words: galaxies: distances and redshifts – galaxies: photometry – galaxies: spiral – cosmology: observations – large-scale structure of Universe – infrared: galaxies.

1 INTRODUCTION

The Milky Way obscures the background extragalactic sky through dust extinction, notably within $\pm 5^\circ$ of the Galactic Plane, and stellar confusion. This results in the so-called Zone of Avoidance (ZOA), literally meaning that the region is devoid of galaxies, and accordingly, extragalactic astronomers avoid working in this region of the sky. This has led to a gap in our understanding of large-scale structure (LSS) in the ZOA, especially the conspicuous features

that cross the ZOA, such as the Perseus-Pisces Supercluster (PPS; Einasto, Joeveer & Saar 1980; Giovanelli & Haynes 1982; Focardi, Marano & Vettolani 1984; Hauschildt 1987), the Great Attractor (GA; Dressler et al. 1987; Lynden-Bell et al. 1988; Woudt, Kraan-Korteweg & Fairall 1999) and the Local Void (LV; Tully & Fisher 1987; Kraan-Korteweg et al. 2008).

Galaxy peculiar velocities, i.e. deviations from the isotropic Hubble expansion, are sensitive to LSS such as the GA (Dressler et al. 1987) and can be used alongside redshift surveys for cosmography and to study the dynamics of the Local Group (LG), cosmic flow fields, and the origin of the observed dipole in the cosmic microwave background (CMB) (e.g. Masters et al. 2006; Springob et al. 2007;

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Masters, Springob & Huchra 2008; Magoulas et al. 2012; Courtois et al. 2013; Hong et al. 2014; Mutabazi et al. 2014; Springob et al. 2014; Tully et al. 2014; Carrick et al. 2015; Springob et al. 2016; Tully, Courtois & Sorce 2016). The use of infrared wavelengths in these studies (Masters et al. 2008; Lagattuta et al. 2013; Sorce et al. 2013, 2014; Neill et al. 2014) has minimized the impact of the ZOA, but the most obscured part of the ZOA ($|b| \leq 5^\circ$) remains mostly unexplored.

This project aims to derive the peculiar velocities of galaxies in the ZOA to supplement the above studies and provide a truly all-sky peculiar velocity survey. In this series of papers, we aim to map most of the dynamically important structures in the ZOA such as the GA and LV in greater detail. The requirements to pursue such a project are: (i) a calibrated and unbiased Tully–Fisher (TF) relation to be used as the global template relation; (ii) 21 cm observations of spiral galaxies in the ZOA from which to extract the redshift and the rotational velocity of galaxies; (iii) follow-up NIR imaging of the H I sources to measure the apparent magnitude of each galaxy. With these three ingredients, the absolute magnitude can be derived using the template relation given the rotational velocity. The distance to each galaxy can then be calculated independent of the redshift using the distance modulus.

The TF template is constructed by Said, Kraan-Korteweg & Jarrett (2015) using the isophotal magnitudes of the 888 spiral galaxies from Masters et al. (2008). The advantage of using isophotal and not total magnitudes is twofold. First, they are more robust and can be measured easily both in and out of the ZOA. Secondly, they are consistent between different surveys which means they can be used to combine data from different data sets without corrections (see fig. 1 in Said et al. 2015). A correction model for the change in the shape of galaxies due to dust extinction is also presented in Said et al. (2015). Said et al. (2016) presents the second element for this ZOA peculiar velocity survey which is the 21 cm H I-line spectra of inclined, $(b/a)^\circ < 0.5$, spiral galaxies. The average signal-to-noise ratio for this H I survey was 14.7 which is adequate for TF studies. Five different types of line-widths are presented to select the most robust one. Conversions between these widths are derived to allow combination of data from different surveys. This third paper is dedicated to systematic NIR follow-up observations of all galaxies in the HIZOA catalogue (HIZOA-S; Staveley-Smith et al. 2016, HIZOA-N; Donley et al. 2005, GB; Kraan-Korteweg et al. 2008).

At low Galactic latitude NIR wavelengths are preferred over optical wavelengths because of the ability of NIR radiation to penetrate through dust. The NIR also provides a more stable indicator of total stellar mass (Kraan-Korteweg & Lahav 2000; Kraan-Korteweg 2005). In the last few decades many surveys have used the NIR to unveil the LSS hidden behind the Milky Way. The 2MASS extended source catalogue (2MASX) contains galaxies that have never been seen before in the ZOA (Jarrett et al. 2000). Jarrett et al. (2000) present new extended sources in the ZOA at Galactic longitude between 40° and 70° . Deeper NIR observations have been used specifically for dedicated surveys in the ZOA. Nagayama et al. (2004) used the same instrument that we have used in this work to conduct an NIR survey around the radio galaxy PKS1343-601. They detected 19 galaxies and another 38 galaxy candidates of which only three were known previously. Woudt et al. (2005) also used the IRSF to obtain deep photometry for 76 galaxies which was used in the determination of the distance to the Norma cluster. Nagayama et al. (2006) used a deep NIR survey of a luminous cluster in the GA region and identified 111 galaxy candidates. Longer wavelengths are also used in the ZOA; Jarrett et al. (2007) used mid-infrared wavelengths to unveil two galaxies in the GA region. A large deep

NIR survey of the Norma Wall (NWS: Riad 2010; Riad, Kraan-Korteweg & Woudt 2010; Kraan-Korteweg et al. 2011; Riad et al., in preparation) was also conducted with the IRSF. This survey resulted in a catalogue of 4360 galaxies with completeness limits of 15.6, 15.3 and 14.8 mag in the J , H , and K_s bands, respectively. Given the success of these surveys, we started a follow-up NIR survey of the HIZOA galaxies in the southern ZOA using the same telescope and instrument (IRSF).

This paper is organized as follows. Section 2 discusses the observations, calibration and observatory site conditions. The final extended source catalogue and parameter characterization are presented in Section 3. Completeness as a function of dust extinction and stellar density is discussed in detail in Section 4. Comparison of the resulting photometry with the 2MASS and UKIDSS surveys is presented in Section 5. We summarize our results in Section 6. All magnitudes are quoted in the Vega System.

2 OBSERVATIONS

Deep NIR follow-up observations of all HIZOA galaxies were conducted with the Infrared Survey Facility (IRSF), a 1.4-m telescope situated at the South African Astronomical Observatory (SAAO) site in Sutherland, South Africa. The Simultaneous three-colour (J , H , and K_s) Infrared Imager for Unbiased Survey (SIRIUS; Nagayama et al. 2003) camera on the IRSF has a field of view of 7.7×7.7 arcmin² (ideally suited for HIZOA follow-up given the 4 arcmin positional accuracy of the H I detections) with a pixel scale of 0.45 arcsec pixel⁻¹ compared to 2.0 arcsec pixel⁻¹ for 2MASS (Skrutskie et al. 2006). A pilot project and the first results of the catalogue were published by Williams, Kraan-Korteweg & Woudt (2014) who presented photometry for 557 galaxies in the HIZOA catalogue with $cz \leq 6000$ km s⁻¹. For completion we have included their 578 fields in the current study.

An additional 105 2MASX fields in the southern ZOA but not in the HIZOA survey were also observed. The H I spectral line data for these 105 2MASX fields are from Parkes observations and available either from the 2MASS TF Survey (Hong et al. 2013) or from Said et al. (in preparation). These additional galaxies are all the bright ($K_s^\circ = 11.25$ mag), edge-on ($b/a = 0.5$) 2MASX galaxies in the southern ZOA ($5^\circ \leq |b| \leq 10^\circ$).

2.1 Data acquisition

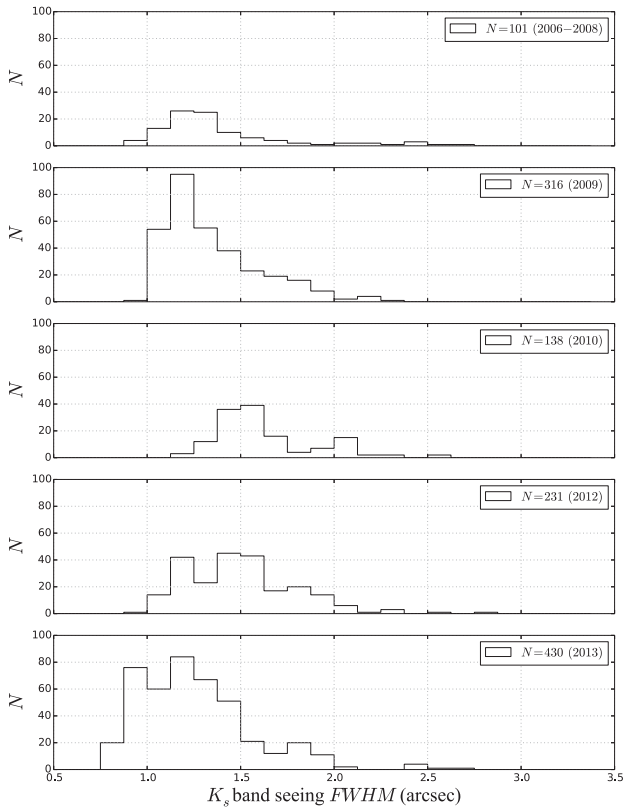
Data acquisition started in 2006 and was completed by 2013, resulting in deep NIR imaging of all the HIZOA targets. The images have exposure times of 10 min, and are 2 mag deeper than 2MASS in the K_s -band (Riad 2010). We used the dithering technique to overcome the problem of faulty pixels in the NIR detector. We repeated a 24 s exposure 25 times with a dithering step of 15 arcsec. This dithering resulted in increasing the final image size to 8.6×8.6 arcmin². A total of 12 weeks were allocated to this project starting in 2009. 101 fields were observed between 2006 and 2008 as part of other projects to test the feasibility of this project. Table 1 shows the observations, number of allocated weeks, number of fields observed and the Observer In Charge (OIC). Substantial time was lost during the 2009 and 2010 runs due to bad weather and cooling system problems.

2.2 Data reduction and calibration

The primary data reduction, including dark frame subtraction, flat-field correction, sky-subtraction, dither combination and

Table 1. IRSF observations and the Observer In Charge (OIC). OIC: Wendy Williams (WW), Tom Mutabazi (TM) and Khaled Said (KS).

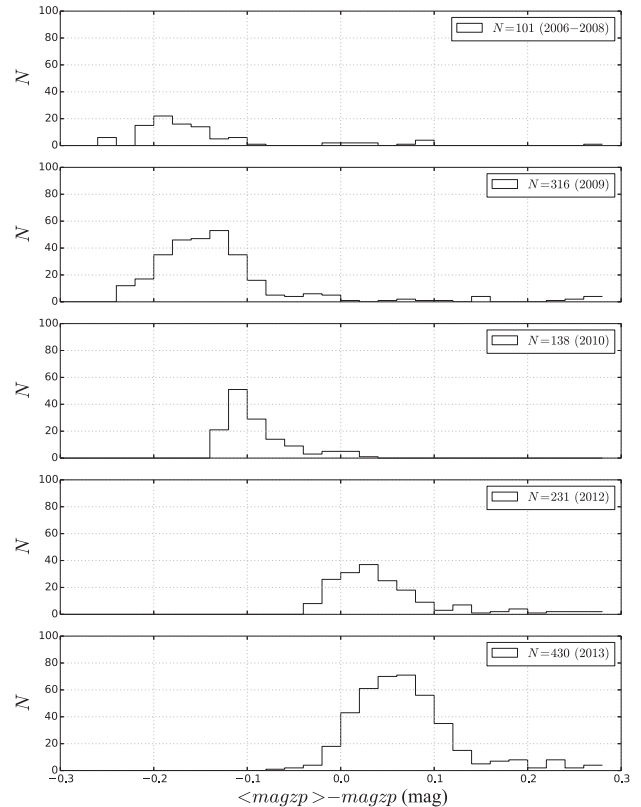
Year	Month(s)	Allocated weeks	No. of fields	OIC
2006–2008			101	
2009	March/April	2	249	WW
2009	March/April	2	67	WW
2010	June/July	3	138	WW
2012	May	2	231	TM
2013	April	3	430	KS

**Figure 1.** Distributions of the measured K_s -band seeing FWHM from 2006 to 2013.

astrometric and photometric calibration, was carried out using the pipeline software for SIRIUS.¹

2.3 Observatory site conditions and quality control

Figs 1 and 2 show the distributions of the K_s -band seeing and the deviation of magnitude zero-point from the mean, respectively, for the 1229 observed fields as a function of observation date. An average K_s -band seeing of 1.38 arcsec and an average K_s -band magnitude zero-point of 20.1 mag were found for the entire survey. These values agree with those found for the Norma wall survey (Riad 2010; Kraan-Korteweg et al. 2011). Fig. 1 shows that 95 per cent of the survey has seeing values below 2.0 arcsec. Fig. 2 shows that 85 per cent of the survey lies within ± 0.2 mag of the mean magnitude zero-point. While the deviations of magnitude zero-point from

**Figure 2.** Distributions of the K_s -band photometric zero-point magnitude from 2006 to 2013.

the mean are not significant, Fig. 2 shows a clear trend that the sky is getting brighter with time. All fields with either seeing or zero-point magnitude outside of these two ranges were inspected visually to check their image quality. Poor quality fields were re-observed under photometric conditions.²

Fig. 3 shows one field that was re-observed, where the left-hand panel shows the field observed under non-photometric conditions and the right-hand panel shows that same field re-observed under photometric conditions. The image on the left-hand side was taken under a partially cloudy sky which affects both the seeing and magnitude zero-point. The K_s -band seeing for the left-hand side image is 2.6 arcsec. The K_s -band zero-point magnitude for the image on the left-hand side was found to be 18.1 mag. In contrast, the right-hand side image shows the field observed under photometric conditions. The K_s -band seeing and zero-point magnitude for the right-hand side image are 0.9 arcsec and 20.1 mag, respectively. The counterpart of the HIZOA galaxy, J1624-45A, in this field is an edge-on galaxy and is marked with the white ellipse in the bottom-left corner of each image. This example demonstrates how important photometric conditions are for this kind of survey. The extent of the galaxy on the left-hand side is underestimated while the real size of the galaxy is apparent in the right-hand image.

² These poor fields were removed and replaced with the photometric-quality ones for the final catalogue but we have made both FITS files available for comparison purposes.

¹ <http://irsf-software.appspot.com/yas/nakajima/sirius.html>

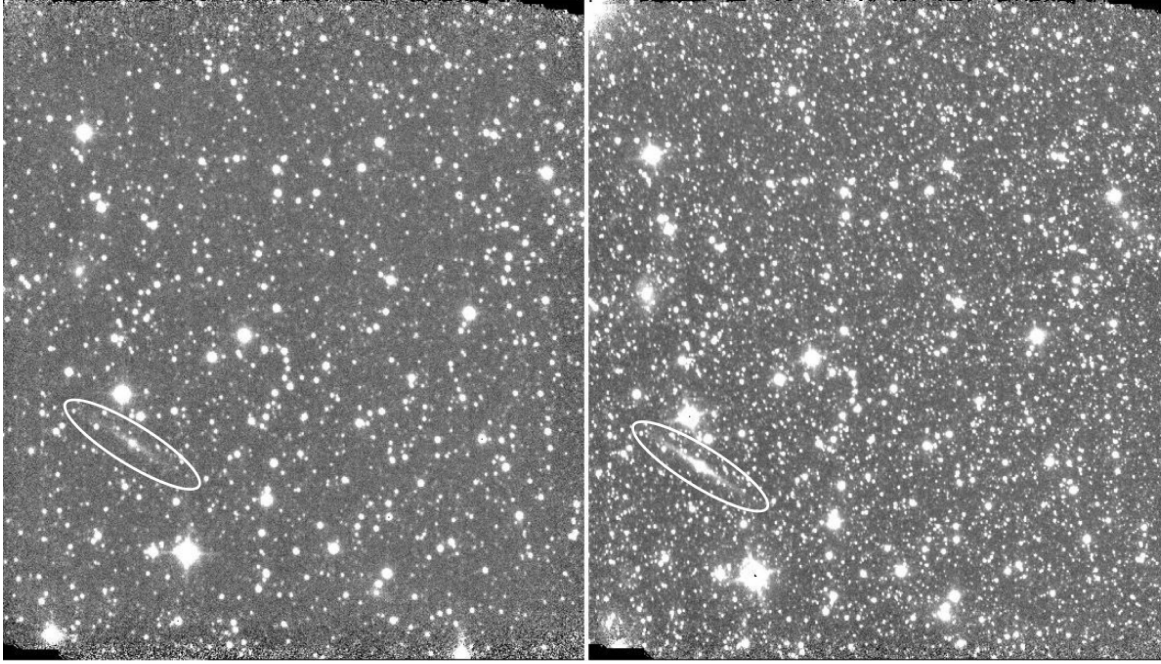


Figure 3. K_s -band photometric versus non-photometric images of the same field. Both images are 8.6×8.6 arcmin; the image on the left was observed under non-photometric conditions while the image on the right was re-observed under photometric conditions. The K_s -band seeing and zero-point magnitude for the image on the left are 2.6 arcsec and 18.1 mag, respectively, while the K_s -band seeing and zero-point magnitude for the image on the right are 0.9 arcsec and 20.1 mag, respectively. Note also the increase in resolved stars in the right image.

3 FINAL EXTENDED SOURCE CATALOGUE

The procedures for source identification, star subtraction and photometry are described in detail by Williams et al. (2014). In this section we will therefore only present the final extended source catalogue. This catalogue contains photometry in the J , H , and K_s bands for 1044 galaxies. We divide the final catalogue into two separate catalogues, one for the HIZOA galaxies and the second for the brightest 2MASX galaxies, full catalogues are available electronically. FITS files for the three bands are also available upon request. An example of the catalogue is given here for the brightest 100 galaxies in the catalogue.

3.1 NIR parametrization and catalogue

The main goal of this paper is to provide accurate NIR photometric parameters for galaxies hidden behind the ZOA. These parameters will be used in the forthcoming NIR TF analysis. In this section we provide the required parameters. For consistency with Williams et al. (2014), we employ the same methods and naming for the whole survey as that described by Cutri et al. (2003). Table 2 and Fig. 4 present the photometry for and the postage stamp images of the brightest 100 galaxies in the catalogue, respectively, in order of isophotal K_{s20} fiducial elliptical aperture magnitude. The catalogue is presented in its entire form online. The NIR parameters listed in the catalogue are as follows.

Column (1) – unique ZOA ID formed from sexagesimal coordinates [ZOAhhmmss.sss±ddmms.ss].

Column (2) – HIZOA ID as reported in the HIZOA survey publications (Donley et al. 2005; Staveley-Smith et al., 2015; Kraan-Korteweg et al., in preparation).³

³ If the galaxy is not in the HIZOA catalogue, the 2MASX ID is provided instead.

Column (3) – Survey name.

Columns (4 and 5) – Right Ascension (RA) and Declination (Dec.) in the J2000.0 epoch.

Columns (6 and 7) – Galactic coordinates [degree].

Column (8) – J -band ellipticity ($\epsilon = 1 - b/a$) measured as the mean value of the ellipticities of the ellipses fitted between the 1σ and 2σ , where σ is the sky rms.

Column (9) – Isophotal K_{s20} fiducial elliptical aperture semi-major axis [arcsec].

Columns (10–12) – J -, H -, and K_s -band K_{s20} fiducial elliptical aperture magnitudes and associated errors [mag].

Columns (13–15) – J -, H -, and K_s -band central surface brightness [mag arcsec⁻²].

Column (16) – $E(B - V)$, Galactic extinction as reported by Schlafly & Finkbeiner (2011).

Column (17) – SD , IRSF stellar density $\log(N_{K_s < 14}/\text{deg}^2)$ where $N_{K_s < 14}$ is the number density of stars brighter than 14 mag in the K_s band.

3.2 Data presentation

Table 3 and Fig. 5 summarize the characteristic properties of the NIR catalogue. In Table 3 we list the mean, maximum and minimum for a number of parameters. These parameters are as follows:

- (i) J -band ellipticity (ϵ_J)
- (ii) K_{s20} fiducial elliptical aperture semi-major axis ($r_{K_{s20}fe}$)
- (iii) J -, H -, and K_s -band K_{s20} fiducial elliptical aperture magnitudes ($J_{K_{s20}fe}$, $H_{K_{s20}fe}$, and $K_{sK_{s20}fe}$)
- (iv) Galactic extinction as reported by Schlafly & Finkbeiner (2011) ($E(B - V)$)
- (v) IRSF stellar density of stars brighter than 14 mag in the K_s band (SD)

Table 2. NIR coordinates, sizes, shapes and photometry of the brightest 100 galaxies in the catalogue in order of isophotal $K_{s,20}$ fiducial elliptical aperture magnitude. The full table is available online.

ZOA ID	HI Name	Survey	RA (J2000)	DEC (J2000)	l [deg]	b	ϵ_J	$r_{K_{s,20}/e}$ [arcsec]	$J_{K_{s,20}/e}$ [mag]	$H_{K_{s,20}/e}$ [mag]	$K_s K_{s,20}/e$ [mag]	μ_{cJ} [mag arcsec $^{-2}$]	μ_{cH}	μ_{cKs}	$E(B - V)$ [mag]	SD
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)	(14)	(15)	(16)	(17)
ZOA141309.873-652020.76	J1413-65	HIZOA	213.291	-65.339	311.326	-3.808	0.44	87.05	7.50 ± 0.02	6.49 ± 0.02	6.09 ± 0.02	12.67	11.69	11.18	1.264	4.374
ZOA151434.147-525921.52	J1514-53	HIZOA	228.642	-52.989	323.594	4.043	0.77	136.09	8.86 ± 0.02	7.91 ± 0.02	7.42 ± 0.02	14.44	13.49	13.01	0.851	4.299
ZOA085728.473-391605.66	J0857-39	HIZOA	134.369	-39.268	261.500	4.100	0.10	45.77	9.18 ± 0.02	8.32 ± 0.02	7.95 ± 0.02	14.09	13.22	12.85	0.619	3.611
ZOA150928.962-523320.67	J1509-52	HIZOA	227.371	-52.556	323.155	4.810	0.79	98.73	9.24 ± 0.02	8.38 ± 0.02	7.97 ± 0.02	14.97	14.05	13.60	0.662	4.090
ZOA081706.065-272720.64	J0817-27	HIZOA	124.275	-27.456	246.973	4.481	0.41	59.84	9.17 ± 0.02	8.41 ± 0.02	8.11 ± 0.02	14.83	14.09	13.81	0.220	3.746
ZOA181427.987-022505.11	J1814-02	HIZOA	273.617	-2.418	26.526	7.099	0.47	47.82	10.04 ± 0.02	8.90 ± 0.02	8.50 ± 0.02	14.72	13.70	13.40	1.980	3.986
ZOA122238.290-583657.66	J1222-58	HIZOA	185.659	-58.616	299.180	4.046	0.69	60.30	9.71 ± 0.02	8.92 ± 0.02	8.60 ± 0.02	14.62	13.83	13.55	0.499	4.121
ZOA151413.880-464827.22	2MASX1514-464	2MASS	228.558	-46.808	326.808	9.338	0.53	114.03	10.75 ± 0.02	10.16 ± 0.02	8.60 ± 0.02	18.23	17.64	16.89	0.217	3.805
ZOA094916.505-475511.27	J0949-47A	HIZOA	147.319	-47.919	274.257	4.549	0.10	38.65	9.67 ± 0.02	8.97 ± 0.02	8.69 ± 0.02	14.37	13.62	13.36	0.301	3.797
ZOA145709.815-542331.46	J1457-54	HIZOA	224.291	-54.392	320.654	4.0957	0.51	53.80	10.04 ± 0.02	9.14 ± 0.02	8.69 ± 0.02	14.96	14.07	13.64	0.731	4.210
ZOA114606.371-562326.95	J1145-56	HIZOA	176.526	-56.391	293.937	5.336	0.27	54.68	9.99 ± 0.02	9.26 ± 0.02	8.82 ± 0.02	16.05	15.26	14.82	0.335	3.844
ZOA135138.534-583515.22	J1351-58	HIZOA	207.910	-58.588	310.724	3.370	0.52	47.27	10.23 ± 0.02	9.31 ± 0.02	8.87 ± 0.02	14.80	13.88	13.41	0.834	4.287
ZOA074752.048-184453.18	J0747-18	HIZOA	116.967	-18.748	236.009	3.374	0.82	84.90	10.06 ± 0.02	9.31 ± 0.02	8.91 ± 0.02	15.79	14.95	14.77	0.327	3.547
ZOA072210.950-055547.38	J0722-05	HIZOA	110.546	-5.930	221.732	4.088	0.78	48.64	9.99 ± 0.02	9.28 ± 0.02	8.99 ± 0.02	14.08	13.41	13.05	0.250	3.528
ZOA070103.346+015439.69	J0701+01	HIZOA	105.264	1.911	212.326	3.011	0.48	41.91	10.16 ± 0.02	9.37 ± 0.02	9.01 ± 0.02	15.09	14.34	13.93	0.520	3.563
ZOA143158.829-552758.82	J1431-55	HIZOA	217.995	-55.466	316.912	4.653	0.10	34.39	10.32 ± 0.02	9.46 ± 0.02	9.05 ± 0.02	15.65	14.75	14.31	0.731	4.091
ZOA080611.134-273140.86	J0806-27	HIZOA	121.546	-27.528	245.709	2.414	0.58	43.68	10.26 ± 0.02	9.51 ± 0.02	9.09 ± 0.02	14.84	14.10	13.67	0.420	3.731
ZOA074843.871-261445.62	J0748-26B	HIZOA	117.183	-26.246	242.586	-0.239	0.32	39.43	10.49 ± 0.02	9.64 ± 0.02	9.10 ± 0.02	15.53	14.56	14.04	0.482	3.723
ZOA163211.878-280530.82	J1632-28	HIZOA	248.050	-28.092	351.084	13.502	0.72	81.48	10.28 ± 0.02	9.45 ± 0.02	9.10 ± 0.02	15.64	14.80	14.46	0.525	3.536
ZOA141036.181-653457.76	J1410-65	HIZOA	212.651	-65.583	310.997	-3.958	0.52	46.18	10.35 ± 0.02	9.48 ± 0.02	9.14 ± 0.02	15.13	14.34	13.98	0.516	4.325
ZOA155524.078-581431.30	J1555-581	2MASS	238.850	-58.242	325.222	-3.570	0.37	65.37	11.32 ± 0.02	9.98 ± 0.02	9.16 ± 0.02	16.32	15.40	14.86	0.641	4.404
ZOA063554.386+110808.32	J0635+11	HIZOA	98.977	11.136	201.262	1.658	0.27	31.91	10.78 ± 0.02	9.63 ± 0.02	9.19 ± 0.02	15.67	14.70	14.29	1.310	3.518
ZOA083439.531-400855.61	J0834-40	HIZOA	128.665	-40.149	259.448	0.122	0.22	34.38	10.02 ± 0.02	9.52 ± 0.02	9.30 ± 0.02	15.30	14.61	14.51	1.823	3.838
ZOA132723.827-572922.23	J1327-57	HIZOA	201.849	-57.490	307.768	5.044	0.74	46.73	10.86 ± 0.02	9.82 ± 0.02	9.30 ± 0.02	15.93	14.75	14.09	0.697	3.980
ZOA085838.795-423157.31	J0858-42	HIZOA	134.662	-42.533	264.125	2.141	0.45	43.00	11.95 ± 0.02	10.10 ± 0.02	9.35 ± 0.02	16.58	14.86	14.05	3.431	3.735
ZOA072456.870-093933.95	J0724-09	HIZOA	111.237	-9.659	225.354	2.942	0.29	32.86	10.48 ± 0.02	9.71 ± 0.02	9.41 ± 0.02	15.54	14.90	14.39	0.340	3.529
ZOA160349.297-605840.50	2MASX1603-605	2MASS	240.955	-60.978	324.235	-6.336	0.81	58.98	10.71 ± 0.02	9.85 ± 0.02	9.45 ± 0.02	15.62	15.02	14.67	0.275	4.036
ZOA094952.868-563235.55	J0949-56	HIZOA	147.470	-56.543	279.808	-2.054	0.24	45.57	11.43 ± 0.02	10.29 ± 0.02	9.58 ± 0.02	16.56	15.33	14.62	1.806	4.159
ZOA154710.092-553240.92	2MASX1417-553	2MASS	214.292	-55.545	314.915	5.320	0.85	58.28	11.27 ± 0.02	10.13 ± 0.02	9.58 ± 0.02	16.54	15.49	15.10	0.535	3.994
ZOA155422.988-612025.58	2MASX1554-612	2MASS	238.596	-61.340	323.130	-5.866	0.56	33.80	10.84 ± 0.02	10.11 ± 0.02	9.58 ± 0.02	15.45	15.09	14.44	0.348	4.042
ZOA141933.320-580850.19	J1419-58B	HIZOA	214.890	-58.147	314.363	2.755	0.59	52.89	11.00 ± 0.02	10.14 ± 0.02	9.59 ± 0.02	16.82	15.82	15.27	1.307	4.363
ZOA163140.118-280606.66	J1631-28	HIZOA	247.917	-28.101	350.997	13.584	0.27	35.29	10.83 ± 0.02	10.07 ± 0.02	9.59 ± 0.02	15.78	14.94	14.48	0.542	3.598
ZOA161710.946-581845.49	2MASX1617-581	2MASS	244.296	-58.313	327.304	-5.542	0.57	31.39	10.83 ± 0.02	10.05 ± 0.02	9.61 ± 0.02	15.11	14.33	13.80	0.260	4.203
ZOA101655.552-485252.32	2MASX1016-485	2MASS	154.231	-48.881	278.525	6.529	0.54	37.76	10.72 ± 0.02	9.74 ± 0.02	9.63 ± 0.02	16.01	15.29	14.91	0.183	3.659
ZOA160425.042-604415.93	2MASX1604-604	2MASS	241.104	-60.738	324.450	-6.205	0.77	40.77	10.73 ± 0.02	10.01 ± 0.02	9.65 ± 0.02	15.08	14.63	14.42	0.257	4.039

Table 2 – continued

ZOA ID	HI Name	Survey	RA (J2000)	DEC (J2000)	l	b	ϵ_J	$r_{K_{20}f_e}$ [arcsec]	$J_{K_{20}f_e}$ [mag]	$H_{K_{20}f_e}$ [mag]	$K_{s,K_{20}f_e}$ [mag]	μ_{cl} [mag arcsec $^{-2}$]	μ_{cl} [mag arcsec $^{-2}$]	$\mu_{cl,KS}$	$E(B-V)$ [mag]	SD
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)	(14)	(15)	(16)	(17)
ZOA090033.110-392626.93	J0900-39	HIZOA	135.138	-39.441	262.020	4.438	0.39	37.83	10.79 ± 0.02	10.15 ± 0.02	9.67 ± 0.02	16.21	15.45	15.04	0.576	3.611
ZOA153603.059-594449.42	2MASXJ1536-5944	2MASS	234.013	-59.747	322.327	-3.226	0.46	41.50	10.92 ± 0.02	10.04 ± 0.02	9.67 ± 0.02	15.87	14.93	14.40	0.833	4.435
ZOA131033.767-580021.18	J1310-57	HIZOA	197.641	-58.006	305.472	4.772	0.27	27.20	10.81 ± 0.02	10.02 ± 0.02	9.68 ± 0.02	15.19	14.52	14.22	0.480	4.080
ZOA161319.695-562349.17	J1613-56	HIZOA	243.332	-56.397	328.261	-3.802	0.35	40.49	10.93 ± 0.02	10.69 ± 0.02	9.68 ± 0.02	16.15	15.39	14.87	0.464	4.425
ZOA074252.007-315959.79	J0742-31	HIZOA	115.717	-32.000	246.932	-4.2257	0.52	26.61	11.15 ± 0.02	10.38 ± 0.02	9.71 ± 0.02	15.39	14.72	14.15	0.714	3.757
ZOA101212.032-623159.40	J1012-62	HIZOA	153.050	-62.533	285.686	-5.121	0.62	39.84	10.84 ± 0.02	10.08 ± 0.02	9.71 ± 0.02	15.80	15.03	14.51	0.249	4.036
ZOA162101.624-360831.49	J1621-36	HIZOA	245.257	-36.142	343.413	9.765	0.55	38.44	10.84 ± 0.02	10.05 ± 0.02	9.71 ± 0.02	15.25	14.44	14.09	0.611	3.789
ZOA070056.215-114734.32	J0700-11	HIZOA	105.234	-11.793	224.511	-3.273	0.70	42.70	11.39 ± 0.02	10.29 ± 0.02	9.72 ± 0.02	16.21	15.12	14.61	0.680	3.595
ZOA072457.474-271516.87	J0724-27	HIZOA	111.240	-27.255	240.882	-5.366	0.36	41.90	11.78 ± 0.02	11.03 ± 0.02	9.72 ± 0.02	16.64	15.89	15.49	0.370	3.668
ZOA181530.143-025348.41	J1815-02	HIZOA	273.876	-2.897	26.220	6.647	0.10	32.73	11.41 ± 0.02	10.26 ± 0.02	9.72 ± 0.02	16.32	15.29	14.62	2.230	3.973
ZOA114948.692-640006.93	J1149-64	HIZOA	177.453	-64.002	296.241	-1.934	0.57	42.66	11.73 ± 0.02	10.42 ± 0.02	9.73 ± 0.02	17.05	15.76	15.05	2.098	4.382
ZOA141710.099-553238.77	J1416-55B	HIZOA	214.292	-55.544	314.916	5.320	0.86	68.61	11.23 ± 0.02	10.10 ± 0.02	9.73 ± 0.02	16.58	15.47	14.96	0.533	3.982
ZOA074141.201-223112.25	J0741-22	HIZOA	115.422	-22.520	238.558	0.239	0.80	48.26	11.31 ± 0.02	10.24 ± 0.02	9.74 ± 0.02	16.54	15.35	14.64	0.568	3.737
ZOA161710.749-581844.59	J1617-58	HIZOA	244.295	-58.313	327.304	-5.542	0.53	33.09	10.82 ± 0.02	10.04 ± 0.02	9.75 ± 0.02	15.19	14.39	14.00	0.258	4.159
ZOA101220.012-471741.58	2MASX1012-471	2MASS	153.083	-47.295	276.983	7.402	0.81	68.92	11.15 ± 0.02	10.21 ± 0.02	9.76 ± 0.02	16.72	15.89	15.26	0.164	3.478
ZOA073008.083-220105.84	J0730-22	HIZOA	112.534	-22.018	236.817	-1.851	0.77	104.61	11.51 ± 0.02	10.04 ± 0.02	9.79 ± 0.02	17.94	16.76	16.38	1.557	3.695
ZOA075220.625-250840.47	J0752-25A	HIZOA	118.086	-25.145	242.052	1.022	0.56	33.38	10.92 ± 0.02	10.17 ± 0.02	9.79 ± 0.02	15.47	14.70	14.35	0.327	3.678
ZOA065010.633-111513.52	J0650-11	HIZOA	102.544	-11.254	222.835	-5.381	0.18	26.46	11.27 ± 0.02	10.36 ± 0.02	9.83 ± 0.02	15.75	14.90	14.42	0.960	3.493
ZOA182226.663-354035.70	J1822-35	HIZOA	275.611	-35.677	357.859	-10.062	0.65	37.90	10.98 ± 0.02	10.14 ± 0.02	9.84 ± 0.02	16.10	15.22	14.79	0.120	4.113
ZOA164634.204-390308.21	J1646-39	HIZOA	251.642	-39.052	344.676	4.067	0.24	39.69	11.94 ± 0.02	10.98 ± 0.02	9.85 ± 0.02	16.67	15.79	15.34	0.990	4.446
ZOA141604.868-651502.53	J1416-65	HIZOA	214.020	-65.250	311.644	-3.821	0.48	36.60	11.23 ± 0.02	10.32 ± 0.02	9.89 ± 0.02	15.77	14.86	14.45	0.662	4.314
ZOA100318.769-645803.19	2MASX1003-645	2MASS	150.828	-64.968	286.322	-7.668	0.58	29.86	10.93 ± 0.02	9.98 ± 0.02	9.90 ± 0.02	15.68	14.93	14.73	0.197	3.759
ZOA105345.693-625013.17	J1053-62	HIZOA	163.440	-62.837	289.956	-2.968	0.67	65.48	11.60 ± 0.02	10.68 ± 0.02	9.90 ± 0.02	17.78	16.80	16.23	0.714	4.234
ZOA080708.583-280309.50	J0807-28	HIZOA	121.786	-28.053	246.266	2.310	0.61	42.80	11.01 ± 0.02	10.29 ± 0.02	9.92 ± 0.02	16.69	16.02	15.56	0.460	3.676
ZOA133732.784-585414.06	J1337-58B	HIZOA	204.387	-58.904	308.867	3.436	0.28	33.07	11.33 ± 0.02	10.31 ± 0.02	9.92 ± 0.02	16.20	15.11	14.48	0.937	4.303
ZOA143927.759-552503.43	J1439-55	HIZOA	219.866	-55.418	317.910	4.281	0.22	30.55	11.17 ± 0.02	10.41 ± 0.02	9.97 ± 0.02	16.75	15.92	15.53	0.550	4.110
ZOA105859.839-502155.66	2MASX1058-501	2MASS	164.749	-50.365	285.236	8.603	0.13	24.61	11.03 ± 0.02	10.11 ± 0.02	9.98 ± 0.02	15.41	14.64	14.33	0.246	3.651
ZOA074901.358-261442.69	J0748-26A	HIZOA	117.256	-26.245	242.618	-0.182	0.27	26.65	11.49 ± 0.02	10.33 ± 0.02	9.99 ± 0.02	15.56	14.61	14.20	0.619	3.837
ZOA113728.729-644822.59	J1137-644	2MASS	174.370	-64.806	295.162	-3.056	0.21	26.51	11.96 ± 0.02	10.63 ± 0.02	10.00 ± 0.02	16.35	15.00	14.46	1.363	4.252
ZOA082837.437-371316.76	J0828-37	HIZOA	227.156	-37.221	256.391	0.904	0.66	36.26	11.42 ± 0.02	10.46 ± 0.02	10.03 ± 0.02	16.18	15.34	14.99	0.920	3.855
ZOA134456.207-654051.40	J1344-65	HIZOA	206.234	-65.681	307.400	-3.380	0.54	34.10	11.38 ± 0.02	10.47 ± 0.02	10.03 ± 0.02	16.09	15.23	14.68	0.870	4.326
ZOA132159.534-543645.62	2MASX1321-543	2MASS	200.498	-54.613	308.390	7.994	0.73	41.15	11.28 ± 0.02	10.45 ± 0.02	10.04 ± 0.02	16.63	15.84	15.52	0.349	3.786
ZOA164421.521-552937.33	J1644-55	HIZOA	251.090	-55.494	331.917	-6.333	0.83	41.24	11.41 ± 0.02	10.41 ± 0.02	10.04 ± 0.02	16.52	15.54	15.03	0.292	4.129
ZOA140835.888-532111.29	2MASX1408-532	2MASS	212.150	-53.353	314.411	7.790	0.33	33.68	11.00 ± 0.02	10.35 ± 0.02	10.05 ± 0.02	15.50	15.31	15.22	0.424	3.829
ZOA105842.989-501930.65	2MASX1058-501	2MASS	164.679	-50.325	285.177	8.620	0.82	45.60	11.40 ± 0.02	10.53 ± 0.02	10.08 ± 0.02	16.09	15.27	14.73	0.246	3.651

Table 2 – continued

ZOA ID	HI Name	Survey	RA (J2000)	DEC (J2000)	<i>l</i>	<i>b</i>	ϵ_J	$r_{K_s,20/e}$ [arcsec]	$J_{K_s,20/e}$ [mag]	$H_{K_s,20/e}$ [mag]	$K_s K_s,20/e$ [mag]	μ_{cJ} [mag]	μ_{cH} [mag arcsec ⁻²]	μ_{cKs} [mag]	$E(B - V)$	<i>SD</i>
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)	(14)	(15)	(16)	(17)
ZOA064400.636+122407.13	J0644+12	HIZOA	101.003	12.402	201.040	4.001	0.65	38.29	11.28 ± 0.02	10.55 ± 0.02	10.09 ± 0.02	15.70	15.07	14.46	0.440	3.409
ZOA182423.339-341054.15	J1824-34	HIZOA	276.097	-34.181	359.399	-9.758	0.36	25.92	11.08 ± 0.02	10.38 ± 0.02	10.09 ± 0.02	15.80	15.05	14.72	0.112	4.079
ZOA154526.828-605931.93	J1545-61	HIZOA	236.362	-60.992	322.508	-4.918	0.49	27.22	11.49 ± 0.02	10.58 ± 0.02	10.10 ± 0.02	16.00	14.97	14.41	0.611	4.140
ZOA123157.581-595058.07	J1231-595	2MASS	187.990	-59.850	300.485	2.930	0.28	28.59	11.50 ± 0.02	10.53 ± 0.02	10.11 ± 0.02	16.27	15.32	15.10	0.803	4.215
ZOA151113.653-535743.36	J1511-535	2MASS	227.807	-53.962	322.664	3.466	0.12	21.61	11.39 ± 0.02	10.52 ± 0.02	10.12 ± 0.02	15.70	15.03	14.87	0.858	4.316
ZOA085809.386-454812.51	J0858-45A	HIZOA	134.539	-45.803	266.544	-0.062	0.30	28.25	12.14 ± 0.02	10.95 ± 0.02	10.12 ± 0.02	17.26	15.87	15.08	2.356	3.890
ZOA085828.676-451630.99	J0858-45B	HIZOA	134.619	-45.275	266.181	0.325	0.29	35.29	12.29 ± 0.02	10.94 ± 0.02	10.12 ± 0.02	17.97	16.54	15.77	3.148	3.819
ZOA133724.550-585221.57	J1337-58B	HIZOA	204.35229	-58.873	308.855	3.469	0.68	34.33	11.48 ± 0.02	10.51 ± 0.02	10.12 ± 0.02	15.54	14.67	14.28	0.937	4.303
ZOA154710.889-590408.56	J1547-59	HIZOA	236.795	-59.069	323.868	-3.538	0.39	47.97	11.32 ± 0.02	10.70 ± 0.02	10.13 ± 0.02	16.82	15.96	15.48	0.550	4.345
ZOA163617.005-421325.00	J1636-421	2MASS	249.071	-42.224	340.998	3.445	0.39	28.61	11.15 ± 0.02	10.54 ± 0.02	10.13 ± 0.02	16.10	15.30	14.72	1.669	4.401
ZOA165408.098-353438.65	J1653-35	HIZOA	253.534	-35.577	348.325	5.120	0.13	18.68	11.47 ± 0.02	10.68 ± 0.02	10.17 ± 0.02	15.53	15.12	14.29	0.930	4.395
ZOA140627.300-575142.26	J1406-57	HIZOA	211.614	-57.862	312.799	3.566	0.27	32.83	11.57 ± 0.02	10.62 ± 0.02	10.18 ± 0.02	16.46	15.50	15.13	0.628	4.171
ZOA160441.177-413947.62	J1604-41	HIZOA	241.172	-41.663	337.221	7.994	0.35	28.11	11.25 ± 0.02	10.55 ± 0.02	10.18 ± 0.02	16.32	15.53	15.17	0.559	3.887
ZOA090240.287-413502.67	J0902-413	2MASS	135.668	-41.584	263.902	3.328	0.52	28.06	11.53 ± 0.02	10.61 ± 0.02	10.19 ± 0.02	16.09	15.18	14.80	1.000	3.659
ZOA155335.142-614059.08	2MASX1553-614	2MASS	238.396	-61.683	322.836	-6.068	0.74	38.79	11.32 ± 0.02	10.51 ± 0.02	10.20 ± 0.02	16.10	15.50	15.47	0.287	4.020
ZOA182700.997-203159.00	J1826-20	HIZOA	276.754	-20.533	11.924	-4.088	0.70	41.25	11.59 ± 0.02	10.62 ± 0.02	10.20 ± 0.02	16.62	15.62	15.17	0.714	4.688
ZOA140621.248-602544.76	J1406-602	2MASS	211.589	-60.429	312.051	1.109	0.42	22.18	12.47 ± 0.02	10.97 ± 0.02	10.21 ± 0.02	16.55	15.26	14.53	2.962	4.638
ZOA134423.990-522211.32	2MASX1344-522	2MASS	206.100	-52.370	311.087	9.661	0.75	40.16	11.57 ± 0.02	10.64 ± 0.02	10.22 ± 0.02	16.65	15.79	15.43	0.387	3.632
ZOA170643.848-482357.97	2MASX1706-4823	2MASS	256.683	-48.399	339.642	-4.607	0.71	28.00	11.68 ± 0.02	10.79 ± 0.02	10.24 ± 0.02	16.10	15.27	14.60	0.747	4.363
ZOA080953.826-414136.58	J0809-41	HIZOA	122.474	-41.694	258.049	-4.600	0.84	66.41	12.03 ± 0.02	10.82 ± 0.02	10.25 ± 0.02	17.66	16.41	15.72	1.049	3.750
ZOA151548.734-600409.37	J1515-60B	HIZOA	228.953	-60.069	320.037	-2.083	0.17	32.38	12.09 ± 0.02	10.84 ± 0.02	10.25 ± 0.02	17.56	16.25	15.61	3.113	4.611
ZOA165805.966-211622.32	J1658-21A	HIZOA	254.525	-21.273	0.331	13.190	0.12	28.09	11.30 ± 0.02	10.52 ± 0.02	10.27 ± 0.02	16.21	15.52	15.39	0.300	3.667
ZOA105838.725-645044.49	J1058-645	HIZOA	164.661	-64.846	291.308	-4.551	0.16	22.02	11.53 ± 0.02	10.65 ± 0.02	10.28 ± 0.02	15.82	14.93	14.84	0.533	4.122
ZOA160449.497-414301.20	J1604-41	HIZOA	241.206	-41.717	337.204	7.937	0.44	36.21	11.42 ± 0.02	10.71 ± 0.02	10.29 ± 0.02	16.40	15.62	15.24	0.559	3.887
ZOA183155.989-314742.59	J1831-31	HIZOA	277.983	-31.795	2.280	-10.134	0.35	32.63	11.24 ± 0.02	10.57 ± 0.02	10.30 ± 0.02	16.16	15.49	15.21	0.146	4.054
ZOA063556.737+143557.75	J0635+14B	HIZOA	98.986	14.599	198.185	3.254	0.20	21.72	11.49 ± 0.02	10.66 ± 0.02	10.31 ± 0.02	15.74	14.92	14.65	0.570	3.529
ZOA120920.790-622912.31	J1209-62	HIZOA	182.337	-62.487	298.092	-0.011	0.36	22.97	13.08 ± 0.02	11.21 ± 0.02	10.31 ± 0.02	17.35	15.59	14.70	3.294	4.553
ZOA141232.785-563433.93	J1412-56A	HIZOA	213.137	-56.576	313.972	4.547	0.46	42.52	11.43 ± 0.02	10.62 ± 0.02	10.32 ± 0.02	17.21	16.41	16.07	0.559	4.149
ZOA072653.624-073252.01	J0726-07	HIZOA	111.723	-7.548	223.715	4.362	0.21	27.41	11.00 ± 0.02	10.64 ± 0.02	10.35 ± 0.02	16.59	16.08	15.80	0.240	3.492
ZOA100655.897-450248.59	2MASX1006-450	2MASS	151.733	-45.047	274.892	8.682	0.65	33.97	11.43 ± 0.02	10.72 ± 0.02	10.35 ± 0.02	16.53	15.93	15.52	0.132	3.504

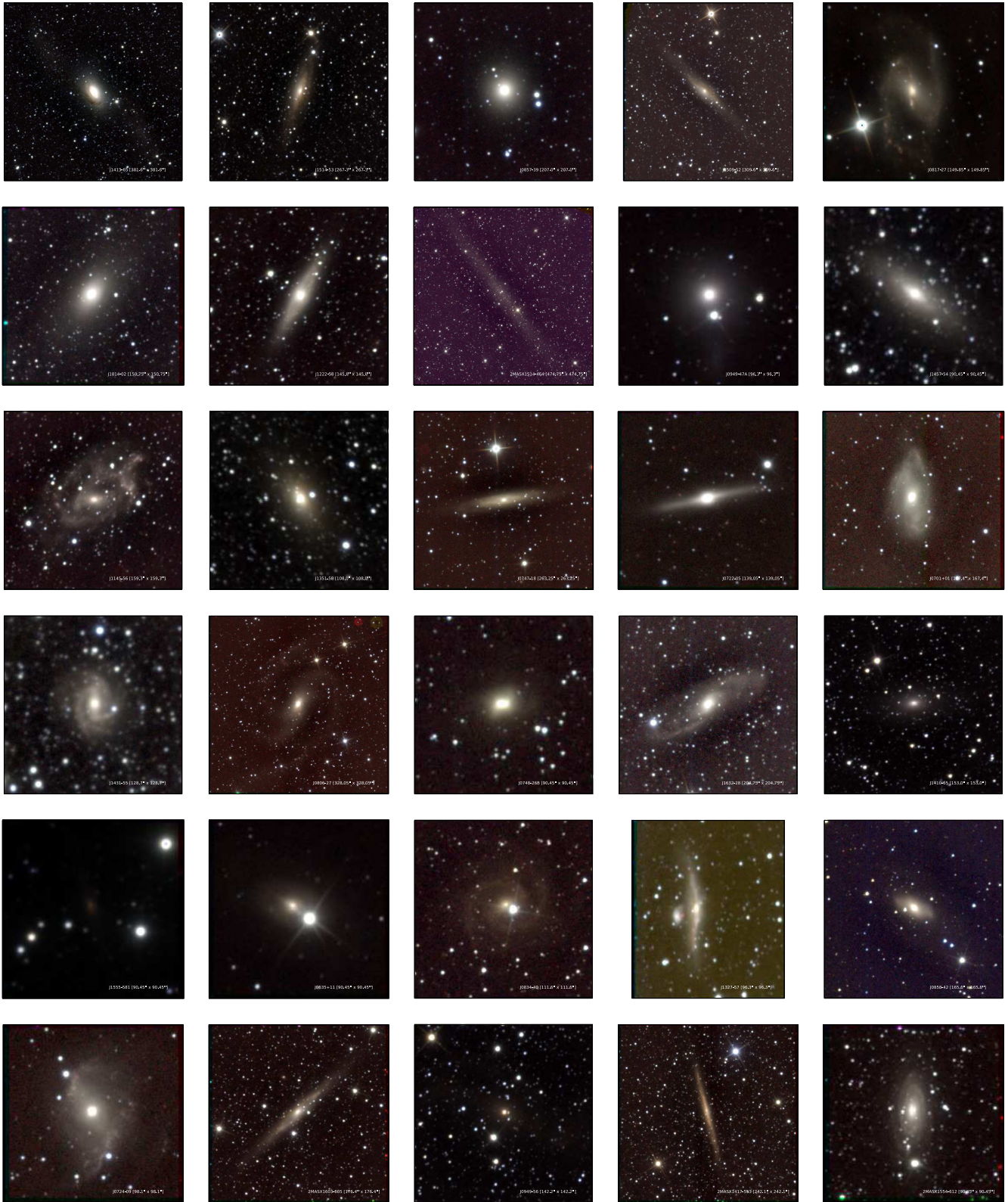


Figure 4. Postage stamp images of the brightest 100 galaxies in the catalogue in order of isophotal K_{s20} fiducial elliptical aperture magnitude. The colour composites are generated in the standard fashion: blue – J band, green – H band, and red – K_s band.



Figure 4 – *continued*



Figure 4 – continued

Figure 4 – *continued***Table 3.** Summary of the characteristic properties of the catalogue.

Parameter	Mean	Max.	Min.
ϵ_J	0.42	0.90	0.10
r_{K_s20fe} [arcsec]	15.57	136.08	0.51
J_{K_s20fe} [mag]	14.28	22.08	7.50
H_{K_s20fe} [mag]	13.39	22.65	6.49
K_{sK_s20fe} [mag]	13.02	23.02	6.09
$E(B - V)$ [mag]	0.92	12.76	0.09
SD	3.95	5.32	3.28

Fig. 5 shows histograms of all parameters except $E(B - V)$ and stellar density which are discussed in detail in the next section as part of the completeness. The top panels, A, B, and C, of Fig. 5 show the distributions in Galactic longitude, shape and size. Panel A illustrates the success of this work in unveiling galaxies hidden behind the MW especially in over-dense regions like Puppis and the GA around $240^\circ < l < 270^\circ$ and $300^\circ < l < 330^\circ$, respectively. The drop in the number of sources towards the Galactic bulge is due to the LV (Kraan-Korteweg et al. 2008; Staveley-Smith et al. 2016). Panel B shows a fairly flat distribution of galaxy ellipticities, which is consistent with the expectation of a random sample of disc galaxies. However, for our final TF sample we use only edge-on galaxies after applying the axial ratio correction from Said et al. (2015). In panel C we plot a histogram of the distribution of the K_{s20} fiducial elliptical aperture semi-major axis. Panel C shows that only three galaxies in our sample have r_{K_s20fe} larger than 100 arcsec and 21 galaxies are larger than 50 arcsec. The largest three galaxies are J1514-53, 2MASX1514-464, and J0730-22 which have r_{K_s20fe} of 136, 114 and 104 arcsec, respectively.

In panels D, E, and F of Fig. 5, we show histograms of the J -, H -, and K_s -band K_{s20} fiducial elliptical aperture magnitudes. Based on the deep NIR survey of the Norma Wall (Riad 2010; Kraan-Korteweg et al. 2011), which used the same instrument with the same setup, we expect similar completeness limits of $J^\circ = 15.6$, $H^\circ = 15.3$, and $K_s^\circ = 14.8$ mag in the J , H , and K_s bands, respectively. However, these limits are only valid for $A_{K_s} < 1.0$ mag and $\log(N_{K_s < 14}/\text{deg}^2) < 4.71$. Compared to other NIR surveys, this IRSF survey is 1 mag deeper than 2MASS in the J band and 2 mag deeper in the K_s band. Moreover, it is only 1 mag shallower than the UKIDSS Galactic Plane Survey (GPS; Lucas et al. 2008)

and VISTA Variables in the Via Lactea (VVV; Amôres et al. 2012) in the K_s band. In Section we present a full comparison with UKIDSS GPS galaxies. The three panels (D, E, and F) of Fig. 5 show that the detection rate drops rapidly for galaxies fainter than 16 mag. However, this survey is not magnitude limited in any sense. Panel F shows that our survey has 63 galaxies brighter than 10 mag.

4 COMPLETENESS

In this section we discuss the completeness as a function of dust extinction, stellar density and $H I$ mass. We divided the catalogue into three sub-samples according to their logarithm $H I$ mass reported by Staveley-Smith et al. (2016). The first column in Fig. 6 shows galaxies with $\log M_{H I} [M_\odot] \geq 9.5$. The second column presents galaxies within the range of $8.5 \leq \log M_{H I} [M_\odot] < 9.5$. The third column shows galaxies with $\log M_{H I} [M_\odot] < 8.5$.

The top panels in Fig. 6 show the completeness as a function of the IRSF stellar density of stars brighter than 14 mag in the K_s band for the three sub-samples. The first two columns in the top panel show that detection rate is 100 per cent for massive galaxies ($\log M_{H I} [M_\odot] > 8.5$) in regions with stellar density of $\log(N_{K_s < 14}/\text{deg}^2) \leq 4$. This detection rate drops to 50 per cent for regions with $\log(N_{K_s < 14}/\text{deg}^2) > 4$. In contrast, the third column in the top panel shows that the detection rate of the least massive galaxies ($\log M_{H I} [M_\odot] < 8.5$) is 75 per cent complete in regions of $\log(N_{K_s < 14}/\text{deg}^2) \leq 4$ and only 30 per cent in regions with stellar density of $\log(N_{K_s < 14}/\text{deg}^2) > 4$. Most of these least massive galaxies are dwarfs and will be excluded from the TF analysis because they have the highest scatter in the TF relation.

Similar conclusions can be drawn from the bottom panels of Fig. 6 which show the completeness as a function of Galactic reddening along the line of sight (Schlafly & Finkbeiner 2011). The first two columns in the bottom panel show that the detection rate of the massive galaxies is nearly 90 per cent in regions of $E(B - V) \leq 1$ mag ($A_V \leq 3.1$ mag). Furthermore, we can still detect massive galaxies up to $E(B - V) = 7$ mag ($A_V = 21.7$ mag). On the contrary, the third column in the bottom panel shows that the detection of the least massive galaxies is not complete anywhere, not even in regions with $E(B - V) \leq 1$ mag ($A_V \leq 3.1$ mag).

We note that the measured photometric parameters (e.g. magnitude, size and shape) will be also affected by these trends and should

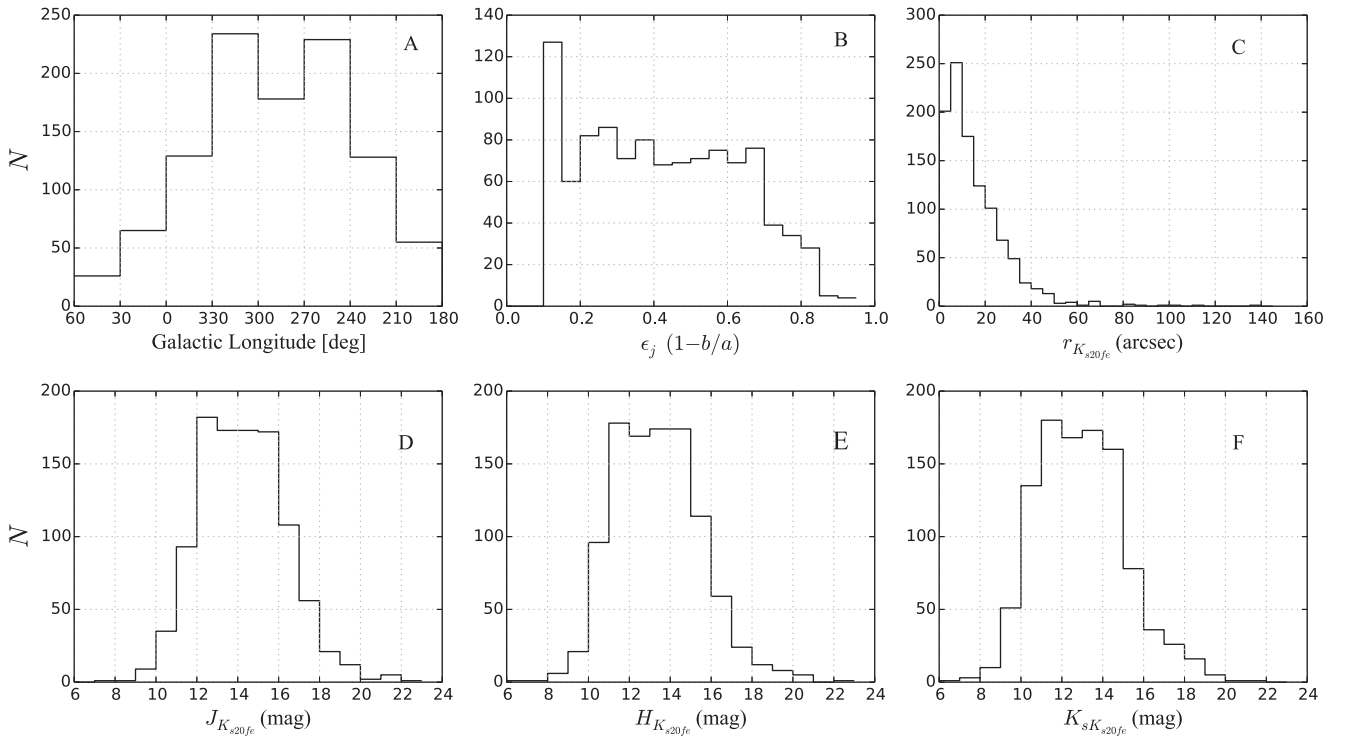


Figure 5. Summary of the characteristic photometric properties of the catalogue. The top panels A, B, and C show the distribution as a function of Galactic longitude, the shape, represented by (ϵ_j) the J -band ellipticity, and size, represented by ($r_{K_{s20fe}}$) K_{s20} fiducial elliptical aperture semi-major axis, of 1044 galaxies in the catalogue. The bottom panels D, E, and F show the distributions of the K_{s20} fiducial elliptical aperture magnitudes in the J , H , and K_s bands, respectively.

be corrected before use in the TF analysis. Riad et al. (2010) discuss the effect of dust extinction on magnitude and size of galaxies and provide a correction model for extinction values up to $A_{K_s} = 3$ mag ($A_V = 25$ mag). While Said et al. (2015) simulate the effect of dust extinction on the shape (ellipticity) of galaxies and also provide a correction model to reproduce the intrinsic axial ratio from the observed value up to extinction levels of $A_J = 3$ mag ($A_V = 11$ mag). These corrections will be used to correct the magnitudes, sizes and shapes of galaxies before use in the TF analysis.

5 COUNTERPARTS AND COMPARISONS

In this section we discuss the counterparts of our survey. We present the confirmed HIZOA counterparts then check for counterparts in the 2MASX (Jarrett et al. 2000) and UKIDSS GPS (Lucas et al. 2008) surveys. We also present a comparison of our photometry with both the shallower 2MASX and the deeper UKIDSS GPS surveys.

5.1 HIZOA counterparts

The pixel size of the final HIZOA cubes of 4 arcmin in RA and DEC makes the IRSF perfect for the follow-up observations given its 8.6×8.6 arcmin² field of view (after dithering). Centering the NIR camera on the HIZOA position has a high probability of locating the counterpart in the image. Thus, the detection of these H I sources depends only on their H I mass and the stellar density and dust extinction of the region in which they lie. Some NIR fields contain more than one possible counterpart to the H I galaxy and

therefore need more attention. All sources identified as H I sources in the NIR fields were inspected by eye and information from their H I profiles used to identify the H I counterpart. A galaxy with double-horn H I profile was normally identified as an edge-on galaxy in the NIR image, while a galaxy with a Gaussian profile was usually identified with a face-on galaxy in the NIR image.

The final NIR catalogue contains counterpart galaxies to 674 sources from all three HIZOA catalogues (HIZOA-S; Staveley-Smith et al. 2016, HIZOA-N; Donley et al. 2005, GB; Kraan-Korteweg et al. 2008). A single counterpart was found for 527 galaxies, while more than one counterpart was found for 147 galaxies.

5.2 2MASS counterparts

Of the 1044 NIR galaxies in the final catalogue, 285 have counterparts in the 2MASX catalogue (180 HIZOA plus 105 2MASS). We used a search radius of only 1 arcsec because of the high positional accuracy of both 2MASS and the IRSF. In Fig. 7 we compare our measured K_s -band K_{s20} fiducial elliptical aperture magnitudes with the same parameter from the 2MASX catalogue (Jarrett et al. 2000, 2003) for these 285 galaxies.

Fig. 7 shows good agreement between the K_{s20} measured for this catalogue and the same parameter reported in the 2MASX catalogue. Small systematic deviations are visible for both faint galaxies (>14 mag) marked as blue circles and bright galaxies (<9 mag) marked as red circles. The deviation for bright galaxies is due to the difference of the pixel scale between these two instruments. The IRSF has pixel scale of 0.45 arcsec pixel⁻¹ compared to

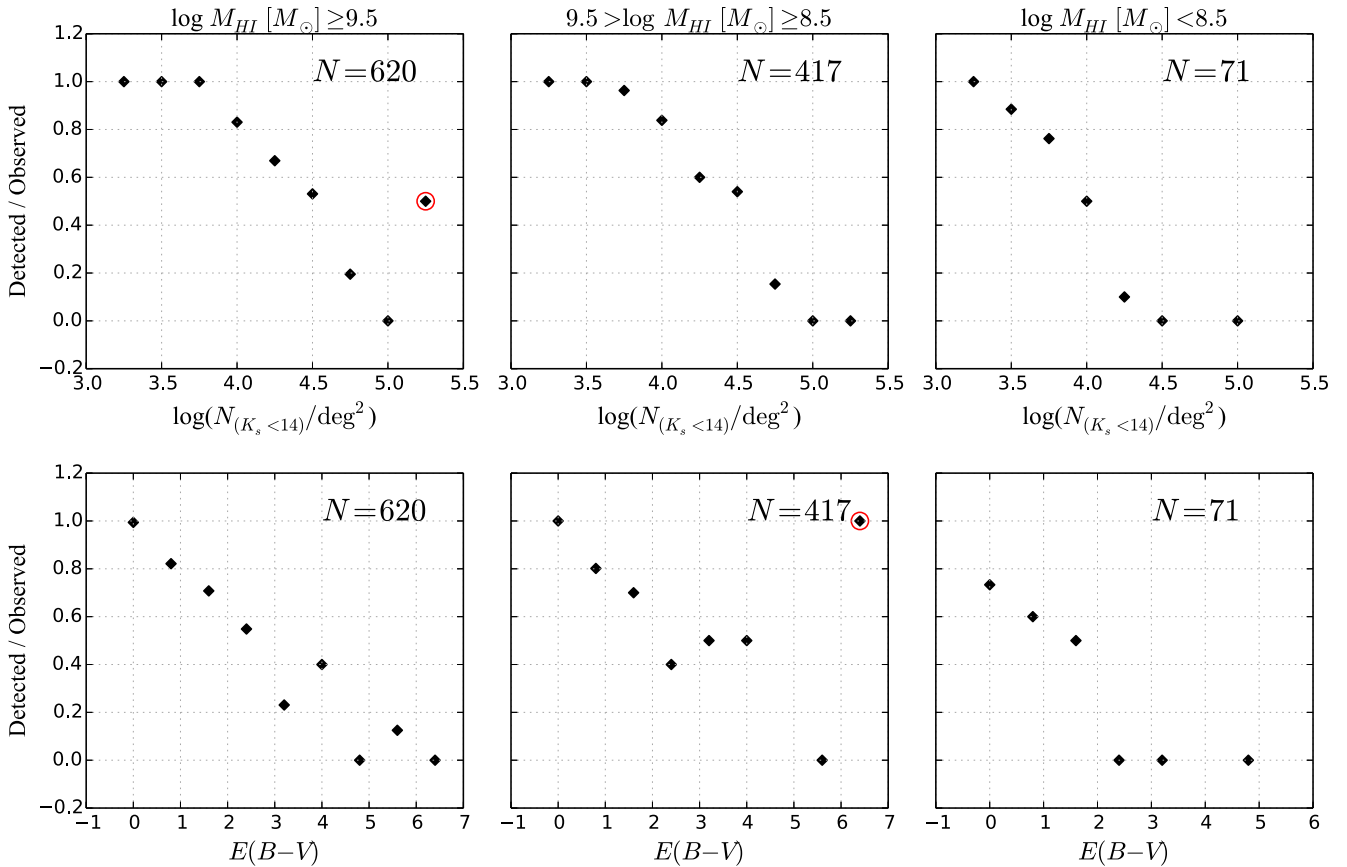


Figure 6. The completeness as a function of stellar density and dust extinction for different H I mass ranges. The three columns present different H I mass ranges. The top panels show the completeness as a function of the logarithm of the stellar density of stars brighter than 14 mag in the K_s band. The bottom panels show the completeness as a function of Galactic reddening along the line of sight. The outliers marked with the red circles are due to low numbers of galaxies in these two bins.

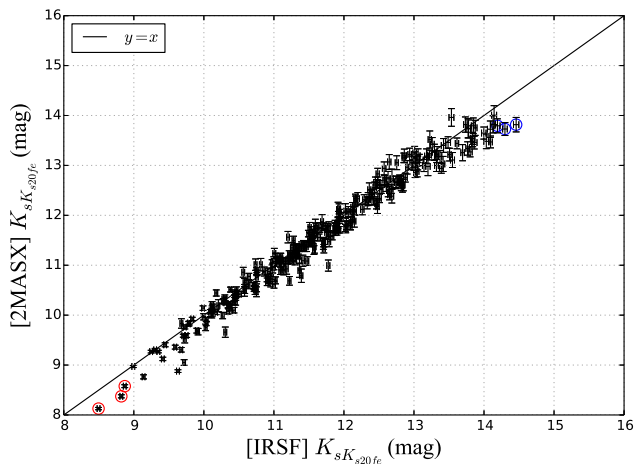


Figure 7. A comparison between the K_s -band K_{s20} fiducial elliptical aperture magnitudes in this catalogue and the same parameter reported in the 2MASX catalogue.

2.0 arcsec pixel⁻¹ for the 2MASX survey. Small, faint stars superimposed on bright galaxies can not be resolved by 2MASX. Thus, the magnitudes of these galaxies are over-estimated because these stars are not subtracted from the image before measuring the photometry. The high resolution of the IRSF instrument leads to more

effective star-subtraction and thus more accurate photometry which is vital when working in the ZOA. The deviation for faint galaxies is a bias due to the completeness limit of 2MASX. The scatter should to be on both sides of the one-to-one line but there are no 2MASX galaxies fainter than 14 mag.

5.3 UKIDSS counterparts

The UKIDSS GPS (Lucas et al. 2008) overlaps with the HIZOA survey in the northern extension published by Donley et al. (2005). We used the publicly accessible *UKIDSS DR8 plus* data release to search for counterparts. Given the high positional accuracy of both the IRSF and UKIDSS GPS, the minimum available search radius of 3 arcsec was used. We found 30 confirmed counterparts in the UKIDSS GPS survey. A modified version of our IRSF photometry pipeline was used to consistently measure the photometric parameters for these galaxies from the UKIDSS GPS images.

Fig. 8 shows the comparison between the K_s -band K_{s20} fiducial elliptical aperture magnitudes measured from the IRSF images and the same parameter measured from the UKIDSS GPS images. The solid line in Fig. 8 is the one-to-one line. Excellent agreement between the IRSF and UKIDSS GPS photometry can be seen. The pixel scale for the UKIDSS GPS images is 0.4 arcsec pixel⁻¹, which is comparable to the IRSF pixel scale of 0.45 arcsec pixel⁻¹. The average seeing of 1.54 arcsec for the UKIDSS GPS images used in this comparison is similar to that of 1.38 arcsec for the IRSF

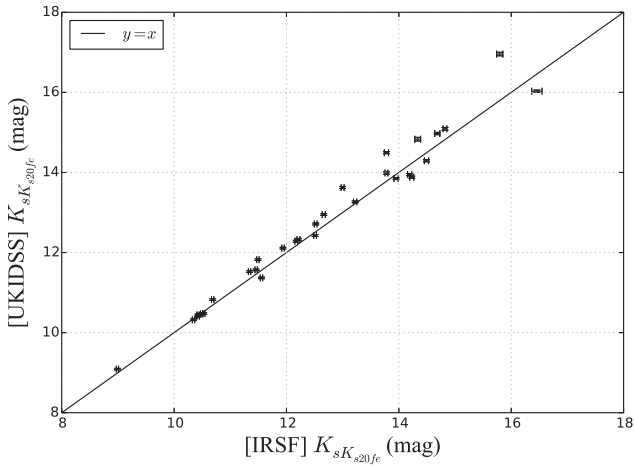


Figure 8. A comparison between the K_s -band K_{s20} fiducial elliptical aperture magnitudes measured from the IRSF images and the same parameter measured from the UKIDSS GPS images using a modified version of the IRSF photometry pipeline. The solid line is the one-to-one line.

survey. This agreement shows that the IRSF imaging does not suffer from foreground contamination, even after star removal, nor does it adversely underestimate the isophotal flux of the ZOA galaxies. It is satisfying to see this agreement between photometric parameters because we plan to extend the current TF project to the northern ZOA using the UKIDSS GPS survey along with $H\text{I}$ data from the Nançay Radio Telescope.

6 SUMMARY

In this paper, we present the observations, data reduction and final catalogue for 1044 NIR galaxies, in the J , H , and K_s bands, in the southern ZOA. The observations were conducted between 2006 and 2013 using the IRSF, a 1.4-m telescope situated at the South African Astronomical Observatory site in Sutherland. This resulted in observations of all galaxies in the three blind systematic deep HIZOA surveys (Donley et al. 2005; Kraan-Korteweg et al. 2008; Staveley-Smith et al. 2016).

The quality of the images is discussed in detail. The survey has an average seeing and zero-point magnitude of 1.38 arcsec and 20.1 mag in the K_s band, respectively. These values agree well with those from previous surveys done with the same instrument (Riad 2010). The mean error of the measured isophotal magnitudes is 0.02 mag which is sufficient for the TF analysis.

The completeness as a function of stellar density and dust extinction was found to be dependent on the $H\text{I}$ mass of each galaxy. The detection rate was found to be 100 per cent for massive galaxies ($\log M_{H\text{I}}[M_\odot] > 8.5$) in regions with stellar density of $\log(N_{K_s < 14}/\text{deg}^2) \leq 4$. However, for small galaxies ($\log M_{H\text{I}}[M_\odot] < 8.5$) the detection rate is 75 per cent in regions of $\log(N_{K_s < 14}/\text{deg}^2) \leq 4$. This detection rate drops to 50 per cent and 30 per cent for regions within $\log(N_{K_s < 14}/\text{deg}^2) > 4$ for massive and small galaxies, respectively. The same conclusion was found for the detection rate as a function of dust extinction. Although the detection rate was high for massive galaxies in regions up to $E(B - V) = 1$ mag ($A_v = 3.1$ mag), it was very low for small galaxies even in regions with very low dust extinction.

We identified 674 galaxies in the final NIR catalogue that have confirmed counterparts in the three HIZOA catalogues. Counterparts from similar NIR surveys are presented. We found 285

2MASX galaxies have counterparts in our final catalogue. However, only 30 galaxies in the UKIDSS Galactic Plane Survey have counterparts in our final NIR catalogue because UKIDSS GPS only overlaps with our survey in its northern extension.

A comparison of our IRSF pointed observations with the 2MASX and UKIDSS GPS surveys was performed for these galaxies. We found good agreement between the K_s -band K_{s20} fiducial elliptical aperture magnitude presented in this paper and the same parameter reported in the 2MASX catalogue. We detect only small deviations for both faint galaxies (> 14 mag) and bright galaxies (< 9 mag). The deviation for bright galaxies is due to the difference of the pixel scale between IRSF and 2MASX instruments. While the deviation for the faint galaxies is a bias due to the completeness limit of 2MASX.

Good agreement was found between the K_s -band K_{s20} fiducial elliptical aperture magnitudes measured from the IRSF data and the UKIDSS GPS data. This agreement confirms that the IRSF images are of equal quality to the UKIDSS GPS images, which are one magnitude deeper. This indirectly implies that IRSF photometry does not suffer from foreground contamination, after star removal, nor does it underestimate the isophotal flux of the ZOA galaxies. The measurement of UKIDSS photometry is regarded as a pilot project of our TF survey in the northern ZOA.

This paper is the third in a series towards the full ZOA TF analysis. The data presented here will be used with the recently calibrated TF relation in Said et al. (2015) as well as the $H\text{I}$ data presented in Said et al. (2016) to derive distances and peculiar velocities for inclined spiral galaxies in the southern ZOA. An extension of this project into the northern ZOA already started last year with the $H\text{I}$ observations of bright inclined 2MASS galaxies using the Nançay Radio Telescope.

ACKNOWLEDGEMENTS

This work is based upon research supported by the Science Faculty at University of Cape Town (UCT), the South African National Research Foundation and the Department of Science and Technology. Part of this research was conducted by the Australian Research Council Centre of Excellence for All-sky Astrophysics (CAASTRO), through project number CE110001020. WLW acknowledges support from the UK Science and Technology Facilities Council [ST/M001008/1]. We thank Dr Tom Mutabazi and Hiroki Onozato for running some of the observations. We also thank Dr Yasushi Nakajima, Dr Takahiro Nagayama and Dr Ihab Riad for many discussions regarding the reduction pipeline. We are very grateful to Dr Christina Magoulas for her careful reading and comments on the paper. We are grateful to Dr Anja Schroeder for making the 2MASX bright galaxy ZOA list available prior to publication.

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SUPPORTING INFORMATION

Additional Supporting Information may be found in the online version of this article:

Table 2. Summary of the characteristic properties of the catalogue. (<http://www.mnras.oxfordjournals.org/lookup/suppl/doi:10.1093/mnras/stw1887/-/DC1>).

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