

DISSIPATION OF PARALLEL AND OBLIQUE ALFVÉN-CYCLOTRON WAVES—IMPLICATIONS FOR HEATING OF ALPHA PARTICLES IN THE SOLAR WIND

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ABSTRACT

We perform 2.5D hybrid simulations with massless fluid electrons and kinetic particle-in-cell ions to study the temporal evolution of ion temperatures, temperature anisotropies, and velocity distribution functions in relation to the dissipation and turbulent evolution of a broadband spectrum of parallel and obliquely propagating Alfvén-cyclotron waves. The purpose of this paper is to study the relative role of parallel versus oblique Alfvén-cyclotron waves in the observed heating and acceleration of alpha particles in the fast solar wind. We consider collisionless homogeneous multi-species plasma, consisting of isothermal electrons, isotropic protons, and a minor component of drifting α particles in a finite- β fast stream near the Earth. The kinetic ions are modeled by initially isotropic Maxwellian velocity distribution functions, which develop nonthermal features and temperature anisotropies when a broadband spectrum of low-frequency nonresonant, $\omega \leq 0.34 \Omega_p$, Alfvén-cyclotron waves is imposed at the beginning of the simulations. The initial plasma parameter values, such as ion density, temperatures, and relative drift speeds, are supplied by fast solar wind observations made by the *Wind* spacecraft at 1 AU. The imposed broadband wave spectra are left-hand polarized and resemble *Wind* measurements of Alfvénic turbulence in the solar wind. The imposed magnetic field fluctuations for all cases are within the inertial range of the solar wind turbulence and have a Kraichnan-type spectral slope $\alpha = -3/2$. We vary the propagation angle from $\theta = 0^\circ$ to $\theta = 30^\circ$ and $\theta = 60^\circ$, and find that the heating of alpha particles is most efficient for the highly oblique waves propagating at 60° , whereas the protons exhibit perpendicular cooling at all propagation angles.

Key words: acceleration of particles – plasmas – scattering – solar wind – turbulence – waves

1. INTRODUCTION

The dissipation of fluid-scale fluctuations in collisionless plasmas can occur via a turbulent cascade, followed by different kinds of wave–particle interactions. The partitioning of energy between alpha particles, protons, and electrons and the efficiency of the heating depend on the characteristics of the waves, the wavevector direction, and the anisotropy of the fluctuations carrying energy at small scales. Anisotropic fluctuations here refer to different wave power in parallel and perpendicular directions with respect to the background magnetic field. Spacecraft observations of magnetic field fluctuations in the solar wind often provide ambiguous single-point measurements, which makes it hard to uniquely determine the type of the observed fluctuations or their angle of propagation. While parallel Alfvén waves and Alfvénic turbulence are ubiquitous in the solar wind and their existence has been unambiguously proven throughout the heliosphere (Bruno & Carbone 2005), much less is known about obliquely propagating waves in the interplanetary medium. The most common electromagnetic fluctuations, which can interact with the particles at the ion scales, are parallel and oblique ion-cyclotron waves (ICWs), as well as kinetic Alfvén waves (KAWs). Such plasma waves at kinetic scales are very hard to detect and identify with a single spacecraft as their interpretation relies on preliminary assumptions. Yet, recent analysis of solar wind data from *Helios*, *MESSENGER*, *STEREO*, *Wind*, and *Cluster* spacecraft show compelling evidence for the existence of parallel and oblique ICWs from 0.3 to 1 AU (Jian et al. 2009, 2010, 2014) and KAWs (Sahraoui et al. 2010; Podesta & Gary 2011; He et al. 2012a; Salem et al. 2012) near the Earth. Jian et al. (2009, 2010) have shown that the ICWs

are rather dispersive with a wide range of wavenumbers and frequencies, and a statistically significant fraction of events propagate at small angles with respect to the direction of the ambient magnetic field. Lately, reconstruction of the magnetic helicity based on *Ulysses* and *STEREO/IMPACT/MAG* data from the undisturbed solar wind suggests that the solar wind turbulence at the ion scales can be reconstructed by a superposition of parallel and oblique ICWs/KAWs (Podesta & Gary 2011; He et al. 2012b). What role each of the different types of waves plays in relation to wave–particle interactions, anisotropic heating, and differential acceleration is still an open question, which remains to be solved. The purpose of this paper is to compare the ion heating by parallel and oblique Alfvén-cyclotron waves and determine which propagation angles are most efficient in preferentially heating the alpha particles within the considered low-frequency turbulent wave spectra. Our secondary goal is to examine the evolutionary path of particles in the configuration space set by the ion temperature anisotropies and their plasma β (see Bale et al. 2009; Bourouaine et al. 2011, 2013; Maruca et al. 2012; Wicks et al. 2013). To address these problems, we perform 2.5-dimensional (2.5D) hybrid simulations with fluid electrons, kinetic protons, and a minor component of differentially streaming α particles. We investigate the turbulent evolution, wave scattering, and dissipation of initial broadband spectra of parallel and oblique Alfvén-cyclotron waves in collisionless finite plasma β fast solar wind conditions. The simulation setup is motivated by the magnetic field power spectra and the plasma properties as in situ observed by *Wind* spacecraft near the Earth. We start within the stable middle point where $T_{\perp i}/T_{\parallel i} = 1$ and the plasma $\tilde{\beta}_i = 0.33$ for both ion species; see

definitions in the next section. This point is far from all linear plasma instability thresholds for both protons and the α particles (Maruca et al. 2012; Bourouaine et al. 2013), and yet the protons are cooled while the alpha particles are heated in a perpendicular direction and the parallel temperature of both species increases. Such cooling/heating cannot be described by the Vlasov linear instability theory predictions. It is a nonlinear process, resulting from the turbulent wave-wave couplings and consequent wave-particle interactions with the alpha particles, whose proper treatment requires direct numerical simulations.

The aim of this paper is to provide the first steps toward comparison and understanding of the efficiency of parallel versus oblique small-scale Alfvénic fluctuations for the preferential heating of alpha particles in the fast solar wind. The initial spectrum resembles Alfvénic turbulence within the inertial range with a Kraichnan-type power spectral slope $\alpha = -3/2$. The spectral index for the fluctuations relates to in situ observations and has been calculated based on the Magnetic Field Instrument (MFI) magnetometer data on board the *Wind* spacecraft in the undisturbed solar wind at 1 AU. The electron temperature and the ion properties are obtained, respectively, from the Three-dimensional Plasma and Energetic Particle Investigation (3DP) instrument and the Solar Wind Experiment (SWE) particle detector on board *Wind*. The next section presents the particle data and the magnetic field measurements used in our model to construct a consistent initial state for the hybrid simulations. The description of the 2.5D hybrid model and the theoretical reconstruction of the initial wave spectrum are presented in Sections 3 and 4. The results of the numerical simulations are revealed in Section 5. Discussion and concluding remarks are presented in Section 6.

2. INITIAL STATE FOR THE SIMULATIONS— *Wind* DATA

In order to relate our model to real solar wind conditions, we have constrained the initial state for the simulations to resemble in situ measurements in a finite- β fast stream near the Earth. The initial plasma parameters for the simulations are based on analysis of *Wind* spacecraft observations in the undisturbed fast solar wind at 1 AU. In what follows below we will use the 92 s average of the *Wind*/SWE (Faraday Cup, FC) data to define the orientation of a mean magnetic field, as considered by Kasper et al. (2008, 2013). Thus, we can define parallel and perpendicular components of all quantities with respect to the direction of the mean field. We will refer to the total proton plasma β as the weighed sum of its parallel and perpendicular contributions, $\tilde{\beta}_p = 1/3(\tilde{\beta}_{p\parallel} + 2\tilde{\beta}_{p\perp})$, where the parallel proton plasma β is defined in terms of the parallel component of the proton thermal velocity and the local Alfvén speed, $\tilde{\beta}_{p\parallel} \equiv v_{th,p\parallel}^2/V_A^2$. Analogously, the perpendicular part is given by $\tilde{\beta}_{p\perp} \equiv v_{th,p\perp}^2/V_A^2$. For isotropic plasmas the parallel and perpendicular components are equivalent: $\tilde{\beta}_{p\parallel} = \tilde{\beta}_{p\perp} = \tilde{\beta}_p$. We use *Wind*/MFI and SWE/FC data to select a period with a given proton plasma β , temperature anisotropy, and relative drift speed between the protons and the α particles. Once this interval is selected, we parameterize the magnetic field fluctuations, which are present there. The SWE/FC data samples are taken at 92 s, which sets the lowest temporal resolution from all instruments in our analysis. The SWE instrument (Ogilvie et al. 1995) provides the proton density,

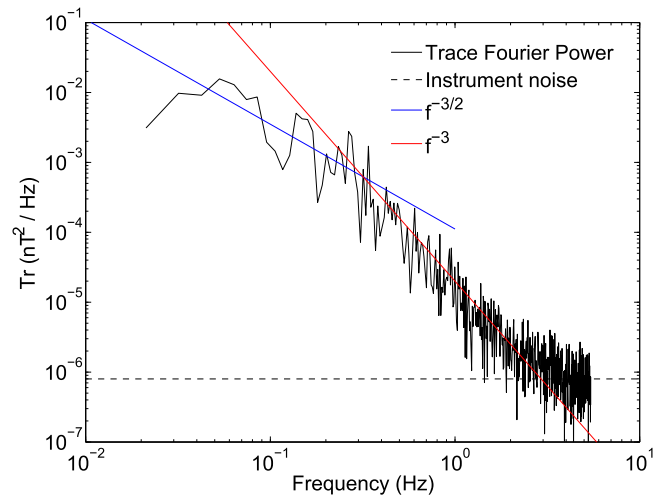


Figure 1. Magnetic field power spectrum based on *Wind*/MFI data fitted with two different power laws. The data are from 2008 January 16, 15:19:09 UT.

temperature, and temperature anisotropy, as well as the α particle density and the relative drift speed between the two ion species. The electron temperature is obtained from the 3DP instrument (Lin et al. 1995) on board *Wind*. MFI data (Lepping et al. 1995) provide vector magnetic field observations every 0.092 s. The different resolution of the instruments therefore leads to 10^3 magnetic fields measurements within the 92 s period of one SWE/FC observation. The analyzed data are from 2008 January 16 at 15:19:09 UT and represent a long-lasting undisturbed fast wind stream during the recent solar minimum. The solar wind speed is 670 km s^{-1} . We have identified the time period with $\tilde{\beta}_p \approx 1$ and $T_{\parallel p}/T_{\perp p} \approx 1$ and selected the data point closest to (1, 1) in the anisotropy- β phase space. Next, we select a scrolling window of 1000 surrounding points of MFI data and Fourier-transform them to obtain spectral information about the magnetic fluctuations. The data analysis has been cross-checked with Morlet wavelet transformations, and the same spectral properties are obtained. The trace of the power spectral tensor, the magnetic helicity, and the variance anisotropy were computed using the local magnetic field direction, following the procedure explained in Wicks et al. (2012). The normalized relative drift speed for the selected interval as obtained from analysis of the data from the SWE/FC particle detector is $V_{\alpha p}/V_A = 0.44$. The proton and the α particle number densities, respectively, are $n_p = 2.65 \text{ cm}^{-3}$ and $n_\alpha = 0.09 \text{ cm}^{-3}$. From the charge neutrality condition the electron number density amounts to $n_e = 2.83 \text{ cm}^{-3}$. The initial temperature for the α particles has not been extracted from the data, and for the rest of the modeling we will assume equal initial temperature for the two ion species. The amplitude of the mean magnetic field as obtained from the 92 s average of the *Wind*/MFI measurements is $|\mathbf{B}_0| \approx 3.85 \text{ nT}$, and the amplitude of the largest magnetic field fluctuation over the time step of 0.092 s is about 0.2 nT. Figure 1 shows the derived power spectral density of the magnetic field fluctuations in units of $\text{nT}^2 \text{ Hz}^{-1}$ as a function of the obtained wave frequency given in Hz (in the spacecraft frame) based on Fourier decomposition of the MFI data. (Morlet wavelet analysis of the data produces the same power spectral slopes.) The level of the instrument noise is overplotted with a dashed line, and the spectra have been fitted by two separate power laws f^α , where α represents the spectral slope.

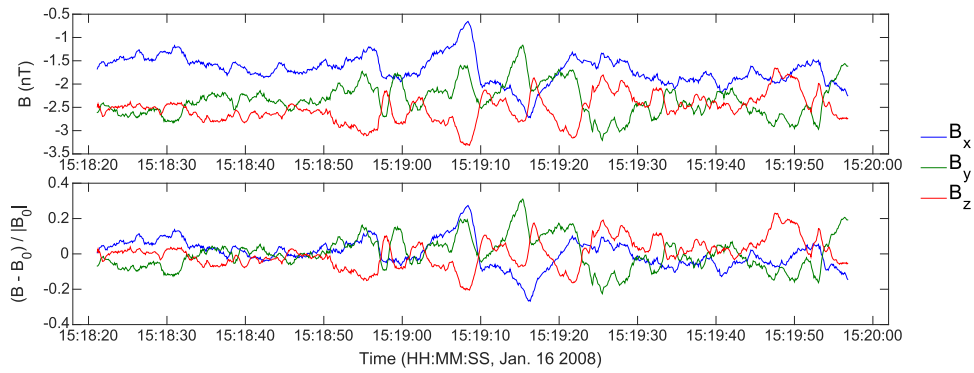


Figure 2. Temporal variation of the magnetic field fluctuations based on *Wind*/MFI data from 2008 January 16. The top panel shows the magnetic field in units of nT, whereas the bottom panel shows the normalized magnetic field components in simulation units with subtracted background magnetic field B_0 .

The noise level of the MFI instrument is not always constant or visible in the spectra, but this particular estimate is in line with the expectations of the *Wind* MFI instrument team (A. Koval 2015, private communication). The original laboratory estimation for the instrument suggests a noise level of $2 \times 10^{-6} \text{ nT}^2 \text{ Hz}^{-1}$. In reality the observed instrument noise floor is often slightly lower than that measured in the laboratory. For this study we have estimated a value of $8 \times 10^{-7} \text{ nT}^2 \text{ Hz}^{-1}$. We should keep in mind that the magnetic field spectra presented in Figure 1 are converted into simulation units (normalized to the background field B_0 and converted to angular frequency). In physical units the instrument noise floor is close to $10^{-4} \text{ nT}^2 \text{ Hz}^{-1}$, in good agreement with the rms noise level of $12 \times 10^{-3} \text{ nT}$ as presented by Lepping et al. (1995). The observed magnetic field fluctuations for the selected time interval are presented in Figure 2. The three components of the magnetic field are normalized to the magnitude of the background field B_0 . The amplitude of the individual fluctuating magnetic field components reaches up to 30% of the background magnetic field magnitude, similar to what has been adopted as an initial state for the simulations (see Figure 6 in the next section). The initial lower-frequency part of the spectrum, associated with large-scale fluctuations, shows an $f^{-3/2}$ power law, typical for the inertial range of a Kraichnan-type MHD turbulence, whereas the higher-frequency (smaller scales) part of the spectrum exhibits a steeper power slope $\alpha = -3$, characteristic for the dissipation range. Similar power slopes have been derived for lower- β data surrounding the (1, 0.33) point in anisotropy- β phase space. In order to study direct cascade processes in the present modeling, we will initialize the simulations with a fraction from the lower-frequency part of the magnetic power spectrum and let the kinetic microturbulence transfer the wave energy down to the ion scales, where the fluctuations can be absorbed as a result of wave-particle interactions.

3. HYBRID CODE—PARTICLES INITIALIZATION

The details of the 2.5D hybrid code and the numerical schemes used in our model have been presented in many previous studies, e.g., Ofman & Viñas (2007), Ofman et al. (2011, 2014), and Maneva et al. (2015). Within the present hybrid model the electrons are described as a charge-neutralizing massless isothermal fluid, and the ions are treated fully kinetically within the particle-in-cell approach. The code allows for two spatial coordinates, but computes all three components of the ion velocities, the current, and the

electromagnetic fields. The electric and the magnetic fields are computed in Fourier space, and periodic boundary conditions are assumed for both particles and fields solvers. The normalization of the temporal and spatial scales is introduced as follows: the simulation time is given in units of inverse proton gyro-frequency Ω_p^{-1} , and the ion velocities are normalized to the Alfvén speed, $V_A \equiv \frac{B_0}{\sqrt{\mu_0 n_e m_p}}$, defined by the electron number density n_e , the proton mass m_p , and the magnitude of the homogeneous background magnetic field B_0 . The length of the simulation box is given in units of the proton inertial length with equal resolution and number of grid points in both special dimensions, so that $L_x = L_y = 384 V_A / \Omega_p$. The simulations are performed with 256×256 cells and 127 particles per cell per species; thus, >16.6 million particles are used. The time step is a small fraction of the proton gyro-period, $\Delta t = 0.02 \Omega_p^{-1}$. The size of each grid cell in x and y direction is the same and equals 1.5 proton inertial lengths. The code uses the expanding box model originally developed by Grappin & Velli (1996). Its application to hybrid models was first presented in Liewer et al. (2001) and has lately been widely used by many authors, e.g., Hellinger et al. (2005), Hellinger & Trávníček (2006, 2013), Ofman et al. (2011, 2014), Moya et al. (2012), and Maneva et al. (2013, 2015). The detailed description of the expanding box model can be found in all of the given references and will not be repeated here. For the purpose of the present study we will only introduce the solar wind expansion parameter $\varepsilon = U_0 / R_0$, defined at $t_0 = 1 / \Omega_p$ for solar wind outflows with constant velocity U_0 at distance R_0 from the Sun. We load the code with homogeneous plasma density and isotropic drifting Maxwellian velocity distribution functions for both protons and the alpha particles. The initial ion densities correspond to the *Wind*/SWE measurements presented in the previous section, which also provide the initial proton temperature, as well as the relative drift speed used to initialize the simulations. The electron temperature is obtained from the *Wind*/3DP instrument. Using the electron number density as a normalization, the normalized number density for the protons and the alpha particles becomes $n_p / n_e \approx 0.94$ and $n_\alpha / n_e \approx 0.03$, respectively. The fluid electron plasma β obtained from the 3DP observations is $\beta_e \equiv 2\mu_0 n_e k_B T_e / B_0^2 = 0.9$. The proton plasma β is defined as the square of the normalized proton thermal speed and is also calculated from the *Wind*/SWE data; see Section 2. We should note that the Alfvén speed used here is defined in terms of the total electron density n_e and the proton mass m_p . In this sense the definition of the ion plasma β for the ion species reads

$\tilde{\beta}_i = 2\mu_0(m_p/m_i)(n_e/n_i)k_B T_i/B_0^2$. The plasma β for the alpha particles is then assumed to be a fraction of the plasma β for the protons, $\tilde{\beta}_\alpha \approx 0.08$; $\tilde{\beta}_p = 0.33$. This selection assures that the two ion species have the same temperature at the beginning of the simulations. The selected *Wind* data interval shows equal parallel and perpendicular temperatures for the protons, and the same is assumed for the α particles. The normalized initial relative drift speed between the two species is $V_{\text{cp}} = 0.44 V_A$. The selected proton plasma-to-cyclotron frequency ratio for the simulations is set to $\omega_{\text{pi}}/\Omega_{\text{ci}} = 0.5 \times 10^4$, as calculated from the proton density and the magnetic field strength provided by the *Wind*/SWE and MFI instruments.

4. CONSTRUCTION OF THE INITIAL TURBULENT WAVE SPECTRA

In the previous section we presented the properties of the hybrid model used for this study and the initialization of the particle properties, such as thermal velocities, temperature anisotropies, and relative drift speeds. This section describes the theoretical reconstruction of an initial Alfvén-cyclotron turbulent wave spectrum, which resembles the *Wind*/MFI data presented in Figure 1. To construct the two-dimensional (2D) oblique wavefronts needed for the 2.5D hybrid simulations, we first generate a one-dimensional (1D) broadband wave spectrum and rotate it in Fourier space to make it 2D, keeping the energy and the magnetic field divergence free. The goal is to start loading the simulations with parallel Alfvén-cyclotron waves, vary the angle of propagation of the magnetic fluctuations, and compare the evolution of the related anisotropic ion heating and differential acceleration for the different cases.

Let us assume that the homogeneous magnetic field background is in the x direction, $\mathbf{B}_0 = B_0 \hat{x}$. The general reconstruction of a 1D broadband spectrum of parallel-propagating magnetic field fluctuations follows the method developed in Viñas et al. (2014) and Vinas et al. (1984):

$$\mathbf{B}(x) = \mathbf{B}_0 + \text{Re} \left[\sum_{k_{\text{min}}}^{k_{\text{max}}} (C_k^L \exp(i\varphi_k)(\hat{y} + i\hat{z}) + C_k^R \exp(i\psi_k)(\hat{y} - i\hat{z})) \exp(ik_{\parallel}x) \right], \quad (1)$$

where k_{min} , k_{max} are the lowest and highest wavenumbers characterizing the wave spectra and

$$\begin{aligned} |C_k^L|^2 &= \frac{\eta^2 B_0^2}{2} \left[\frac{k_{\parallel}^{-\alpha} (1 + \sigma_{mk})}{\sum_{k_{\text{min}}}^{k_{\text{max}}} k_{\parallel}^{-\alpha}} \right] \\ |C_k^R|^2 &= \frac{\eta^2 B_0^2}{2} \left[\frac{k_{\parallel}^{-\alpha} (1 - \sigma_{mk})}{\sum_{k_{\text{min}}}^{k_{\text{max}}} k_{\parallel}^{-\alpha}} \right] \end{aligned} \quad (2)$$

are the Fourier amplitudes of the left-handed and right-handed modes, respectively. The 1D spectrum describes parallel wave propagation, $\mathbf{k} = k_{\parallel} \hat{x}$, with random phases φ_k and ψ_k for each mode. In Equation (2) η is the maximal spectral wave amplitude, $-1 \leq \sigma_{mk} \leq 1$ is the reduced magnetic helicity for each mode, and α is the spectral index of a prescribed power-law power spectral profile. The spectrum conserves the wave

energy, polarization, magnetic helicity, and cross-helicity in configuration space and Fourier space (Viñas et al. 2014).

This general procedure of generating 1D magnetic field fluctuations was applied in a recent 2.5D simulation study (Maneva et al. 2015) to construct 1D initial broadband spectra of left-hand-polarized parallel-propagating Alfvén-cyclotron waves and follow their anisotropic nonlinear turbulent cascade. These simulations assumed that the entire simulation box is permeated by 1D parallel-propagating waves and studied their role for heating and acceleration of alpha particles in the solar wind. In order to initialize our 2.5D numerical simulations with 2D turbulent wave spectra in the present work at $y = 0$, we start with the 1D spectra calculated for all grid points x_i along the direction of the background magnetic field. Next, we transform the resulting 1D field $\mathbf{B}(x)$ into Fourier domain and perform a rotation of angle θ with respect to the z axis to obtain 2D oblique wavevectors. The new wavevectors \mathbf{k}' and the Fourier component of the complex magnetic field $\mathbf{B}'_{\mathbf{k}'}$ acquired through the rotation are given by

$$\mathbf{k}' \equiv k'_{\parallel} \hat{x} + k'_{\perp} \hat{y} = k_{\parallel} \cos(\theta) \hat{x} + k_{\parallel} \sin(\theta) \hat{y} \quad (3)$$

and

$$\mathbf{B}'_{\mathbf{k}'} = -B_{k,y} \sin(\theta) \hat{x} + B_{k,y} \cos(\theta) \hat{y} + B_{k,z}, \quad (4)$$

where

$$\begin{aligned} B_{k,y} &= \frac{1}{2} (C_k^L \exp(i\varphi_k) + C_k^R \exp(i\psi_k)) \\ B_{k,z} &= \frac{i}{2} (C_k^L \exp(i\varphi_k) - C_k^R \exp(i\psi_k)). \end{aligned} \quad (5)$$

We apply this procedure to transform the initial 1D parallel waves into 2D oblique fluctuations at a prescribed angle θ , preserving the energy and the divergence-free property of the initial magnetic field $\nabla \cdot \mathbf{B}' = 0$ (or $\mathbf{k}' \cdot \mathbf{B}' = 0$ in Fourier domain) and obtain an oblique magnetic field Fourier spectrum with power in both parallel and perpendicular directions with respect to the background field. In addition, using the new magnetic spectrum and wavevectors, we construct the bulk velocity spectra for each ion species, namely,

$$\begin{aligned} \delta U_{s,k,y} \pm i U_{s,k,z} &= - \left[\frac{(\omega/k - U_{\parallel s})}{(1 \mp (\omega - kU_{\parallel s})/\Omega_s)} \right] \\ &\times \left(\frac{B'_{k,y} \pm i B'_{k,z}}{B_0} \right), \end{aligned} \quad (6)$$

where s is the species index and the plus (minus) sign represents left(right)-hand-polarized waves.

We repeat the process for each grid row until $y = y_{\text{max}}$, generating N_y oblique spectra all with different random phases. To preserve the total energy initially considered, we must normalize the Fourier coefficients to divide the energy of each perpendicular mode by the number of repetitions made in the k'_{\perp} direction (by construction the energy is well distributed within all the parallel modes k'_{\parallel}). Finally, after the rotation of the entire 1D Fourier spectrum, we obtain oblique 2D spectra of magnetic field fluctuations. Simultaneously, we rotate and compute the corresponding bulk velocities for each ion species as initially given by the parallel dispersion relation, used as our starting case. Applying an inverse 2D fast Fourier transform,

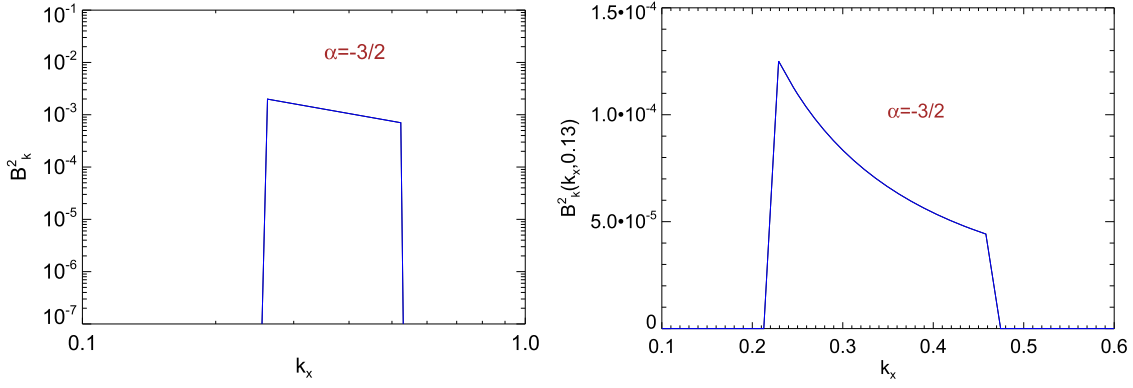


Figure 3. Initial magnetic field power spectra as a function of the parallel wavenumber for the case of parallel (left panel) and oblique (right panel) wave propagation at $\theta = 30^\circ$. The oblique power spectrum is calculated at $k_y = 0.13 \Omega_p/V_A$, but the spectral slope is the same for the entire interval of initial perpendicular wavenumbers $k_y \in [0.13, 0.26] \Omega_p/V_A$.

we retrieve the desired oblique 2D wave spectra of Alfvén-cyclotron waves with both magnetic and velocity field fluctuations, which we load in the 2D spatial grid.

To compute the initial wave frequencies, we solve the warm plasma dispersion relation for parallel wave propagation in a multi-species homogeneous plasma background. We select the left-hand circularly polarized waves and use them to construct the 1D velocity fields. The determinant of the dispersion tensor in the drifting multi-species magnetized plasma is given by the linear Vlasov theory (Davidson & Ogden 1975; Sentman et al. 1981; Gary et al. 2002),

$$\det |\mathbf{D}| = k^2 c^2 - \omega^2 - \sum_s \omega_{ps} \left(\frac{\omega - U_s k}{v_s k} Z(\xi_s^+) + (A_s - 1) \left[1 + \xi_s^+ Z(\xi_s^+) \right] \right). \quad (7)$$

The multi-species index s in the summation stands for electrons, protons, and He^{++} ions. The rest of the notations are as follows: ω_{ps} , v_s , U_s , and A_s are the plasma frequency, the spontaneous velocity, the bulk velocity, and the temperature anisotropy for the s th plasma species, respectively; $A_s = T_{\perp,s}/T_{\parallel,s}$. $Z(\xi)$ is the standard plasma dispersion function with Doppler-shifted frequency-dependent argument

$$Z(\xi_s^+) \equiv \pi^{-1/2} \int_{-\infty}^{\infty} \frac{e^{-x^2}}{x - \xi_s^+} dt, \\ \xi_s^+ = \frac{\omega - U_s k + \Omega_s}{v_s k}.$$

We look for solutions of the above equation for low-frequency waves, with frequencies below the lowest ion-cyclotron in the system, in this case the α -cyclotron. For this study the electron inertia has been neglected and the electron contribution has been eliminated assuming charge neutrality and current conservation. The kinetic plasma dispersion relation is used in order to make sure that the selected initial wave spectrum is stable with respect to Vlasov theory plasma micro-instabilities. Once the frequencies for the parallel waves are obtained, the 1D fluctuations of the transverse velocity fields in the initially isotropic plasma are calculated from the multi-fluid dispersion

relation:

$$\delta \mathbf{V}_{\perp i} = - \frac{\omega - U_i k}{1 - (\omega - U_i k)/\Omega_p} \delta \mathbf{B}_{\perp} / B_0. \quad (8)$$

Then the 1D velocity fluctuations for each species are rotated to obtain the 2D solution for the initial velocity fluctuations in configuration space. The initial wave spectra considered here represent 2D energy-preserving Alfvénic fluctuations and more precisely forward-propagating left-hand-polarized Alfvén-cyclotron waves from the lower-frequency α -cyclotron branch of the dispersion relation. The initial wave spectra demonstrated in Figure 3 are chosen to resemble the end of the inertial range represented by the spectral slope $\alpha = -3/2$ from the *Wind*/MFI spectrum; see Figure 1. The generated spectra consist of 17 modes, and the total amplitude is $\delta B = 0.2 B_0$. For a fixed amplitude and spectral slope, we initialize the simulations with three different propagation angles, 0° , 30° , and 60° , and study their influence on the alpha particles. In the pure parallel case the waves have wavenumbers $k_{\parallel} \in [0.26, 0.52] \Omega_p/V_A$ and frequency range $\omega \in [0.22, 0.34] \Omega_p$. At 30° the spectrum has $k_{\parallel} \in [0.23, 0.46] \Omega_p/V_A$, $\omega \in [0.21, 0.34] \Omega_p$, and $k_{\perp} \in [0.13, 0.26] \Omega_p/V_A$. The highly oblique case at 60° has $k_{\parallel} \in [0.13, 0.26] \Omega_p/V_A$, $\omega \in [0.13, 0.26] \Omega_p$, and $k_{\perp} \in [0.23, 0.46] \Omega_p/V_A$. Figure 3 describes the reconstructed initial magnetic field power spectra as a function of the parallel component of the wavevector. The left panel shows the case of strictly parallel-propagating waves, where $\mathbf{k} = k_x \hat{x}$. The right panel shows the k_x -dependence of the magnetic spectra at fixed perpendicular component $k_y = 0.13 \Omega_p/V_A$ for the case of oblique waves propagating at $\theta = 30^\circ$. The same spectral slope $\alpha = -3/2$ was used in both cases. Since the initial angle is not too large, however, most of the power still remains in the parallel direction and the power law retains its slope. This is no longer true when highly oblique fluctuations at $\theta = 60^\circ$ are considered. Yet, the magnetic field spectrum as a function of the full wavevector remains the same for all angles of propagation and recovers the parallel case shown in the left panel of the figure. We should note that the initial power spectral law for the magnetic field fluctuation is constructed in k -space, whereas the observational data provide information about the power spectra in frequency domain. For the low-frequency MHD part of the

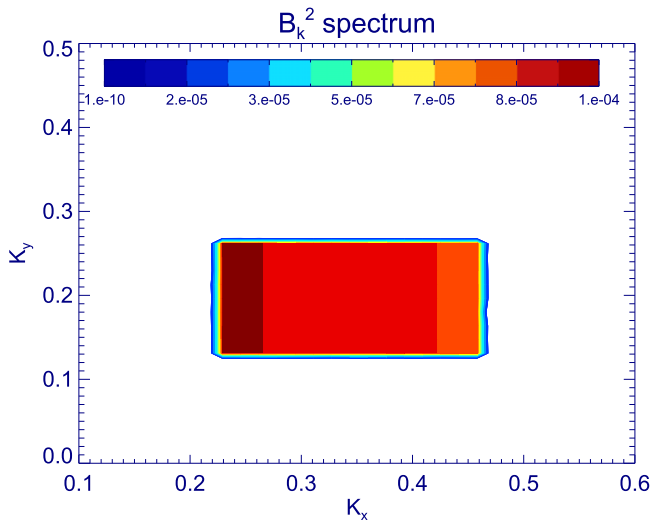


Figure 4. Power spectral density for the initial 2D magnetic field fluctuations as a function of the parallel, k_x , and perpendicular, k_y , wavenumbers for the case of oblique wave propagation at $\theta = 30^\circ$. The magnetic field power is concentrated within the selected wavenumber area presented in Table 1.

Alfvénic spectrum the dispersion relation is linear (waves propagate at the Alfvén speed) and the two descriptions become identical. For the intermediate scales at the beginning of the dispersive part of the spectrum, which we address in our study, the spectral slope of the magnetic field fluctuations in the frequency domain that results from $k_{\parallel}^{-3/2}$ reconstruction will be flatter with a larger corresponding power-law exponent $\alpha > -1$. Figure 4 depicts the 2D power spectral density of the initial magnetic field fluctuations for the case of oblique Alfvén-cyclotron waves propagating at $\theta = 30^\circ$. The figure shows the squared magnitude of the magnetic field fluctuations in Fourier space as a function of the parallel k_x and perpendicular k_y components of the wavevector. The initial wave power is restricted in the wavenumber space described above and presented in Table 1. The following two figures illustrate the corresponding initial magnetic field in configuration space for the case of oblique wave propagation at $\theta = 30^\circ$ and $\theta = 60^\circ$. Figure 5 shows the parallel component of the magnetic field fluctuations as a function of the parallel \hat{x} and perpendicular \hat{y} spatial coordinates, forming the simulation box. The fluctuations are normalized with respect to the magnitude of the background field B_0 , and the spatial coordinates are given in terms of proton inertial length $d_i \equiv c/\omega_{p1} = V_A/\Omega_p$, where c is the speed of light in vacuum and ω_{p1} is the proton plasma frequency. The magnetic field in configuration space for the case of $\theta = 30^\circ$ is given by the inverse Fourier transform of the initial fluctuations, presented in Figure 4. The background field has not been subtracted, which shifts the magnitude of the fluctuations, forcing them to oscillate around the mean normalized value of 1. We should note that in the case of strictly parallel wave propagation there are no fluctuations along the background magnetic field. The nonzero fluctuations in B_x are associated with the presence of oblique waves. Figure 6 is a further illustration of the parallel magnetic field oscillations in the case of oblique wave propagation at $\theta = 30^\circ$ as an example. The two panels represent a cut of the parallel magnetic field component at $y = 32d_i$ (left panel) and

Table 1
Initial Parameters for the Different Numerical Simulation Runs

Case #	θ	$V_{\text{op}} [V_A]$	$\omega_0 [\Omega_p]$	$k_{\parallel} [\Omega_p/V_A]$	$k_{\perp} [\Omega_p/V_A]$
1	0°	0.44	[0.23–0.34]	[0.26, 0.52]	0
2	30°	0.44	[0.21–0.32]	[0.23, 0.46]	[0.13, 0.26]
3	60°	0.44	[0.13–0.26]	[0.13, 0.26]	[0.23, 0.46]

Note. The table presents the simulation cases as a function of the angle of propagation of the initial Alfvén-cyclotron wave spectra, θ , the initial relative drift speed, V_{op} , the range of initial wave frequencies, ω_0 , and parallel (k_{\parallel}) and perpendicular (k_{\perp}) wavenumbers.

$x = 32d_i$ (right panel). At these positions the initial waves are still clearly present in both the parallel direction along x and the perpendicular direction along y regardless of the limited range of selected wavenumbers for the reconstruction of the initial broadband spectra; see Figures 4 and 5.

In the next section, we will present the influence of the initial wave spectra for the anisotropic ion heating for the different propagation angles discussed above. We will discuss the wave scattering and the evolution of the ion velocity distribution functions for protons and α particles in the electromagnetic field of the waves. We will show the evolution of the initial magnetic field and plot the end-stage electromagnetic dispersion relation to analyze the scattering of the initial spectra along with the generation of new wave modes.

5. RESULTS

In this section we will answer the objectives posed in the Introduction, namely, to compare the relative role of parallel and oblique Alfvén-cyclotron waves for the observed preferential heating of alpha particles in the fast solar wind near the Earth. As a secondary goal of this research, we will investigate the evolution of the stable point $\tilde{\beta}_p = 0.33$; $T_{\perp p}/T_{\parallel p} = 1$ in the central domain of the plasma β and temperature anisotropy space for the protons, where the majority of the ion velocity distribution functions are observed (Matteini et al. 2007; Bale et al. 2009; Maruca et al. 2011, 2012; Wicks et al. 2013; Hellinger & Trávníček 2014). This point is stable with respect to linear plasma instabilities, and it is interesting to study its temporal evolution in $(\tilde{\beta}, T_{\perp p}/T_{\parallel p})$ space in the presence of the ambient solar wind turbulence. Figure 7 demonstrates the temporal evolution of the temperature anisotropies for the α particles (top panel), protons (middle panel), and the relative drift speed between the two ion species (bottom panel). The initial plasma state represents the *Wind* measurements explained in Sections 2 and 3. The figure presents a comparison between the different simulation cases (see Table 1) with strictly parallel or oblique initial wave spectra with propagation angle $\theta = 0^\circ, 30^\circ$, and 60° , respectively. The initial state for all three cases is isotropic drifting plasma with $V_{\text{op}} = 0.44 V_A$. The slight initial temperature anisotropy (mainly in the parallel case) is apparent and is due to the nonthermal contribution in the transverse bulk velocity for the ions generated by the initial broadband wave spectra; see Equation (8). The slight initial deviation from the selected value of the relative drift speed in the case of oblique wave propagation is due to the bulk parallel velocity fluctuations induced by the oblique waves. The temperature anisotropy of the α particles throughout the simulations varies with the propagation angle and is highest for the case of parallel wave propagation with final

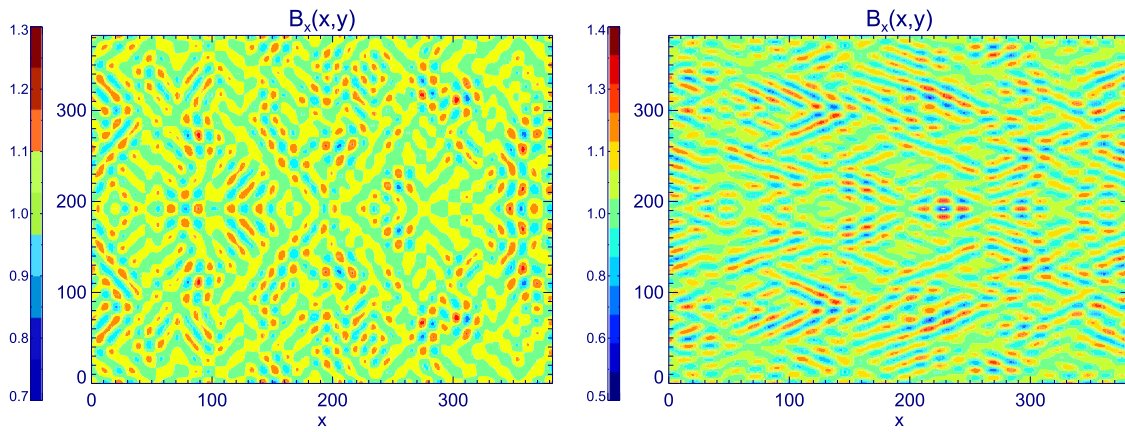


Figure 5. Parallel component of the 2D magnetic field in normalized units in real space. The field corresponds to oblique propagating waves at $\theta = 30^\circ$ (left) and $\theta = 60^\circ$ (right). The spatial coordinates are given in units of proton inertial length. The fluctuations are superimposed on the constant magnetic field background, whose amplitude is used for normalization. The fluctuations introduce inhomogeneity in the uniform background magnetic field.

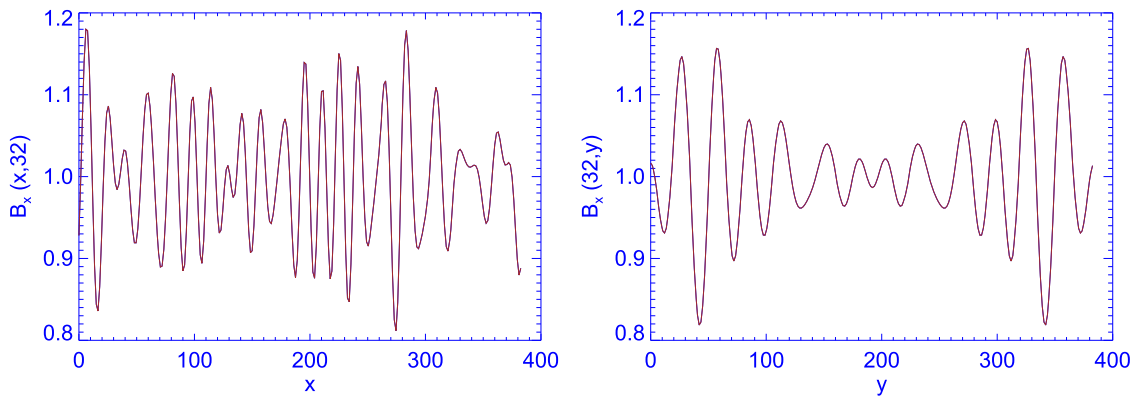


Figure 6. Parallel component of the 2D magnetic field in normalized units as a function of the parallel (left) and perpendicular (right) spatial coordinates, x and y . The fluctuations correspond to obliquely propagating waves at $\theta = 30^\circ$. The plots are complementary to the left panel of Figure 5.

$(T_{\perp}/T_{\parallel})_{\alpha} \approx 1.24$, whereas the anisotropy of the protons smoothly decreases for all propagation angles considered here. Thus, T_{\perp}/T_{\parallel} for the alpha particles increases with 5% for the case of parallel wave propagation, decreases with 5% for $\theta = 30^\circ$, and is reduced by 20% for the highly oblique waves with $\theta = 60^\circ$. The anisotropy of the protons is reduced for all cases: it decreases with 21% in the parallel case, 11% in the case of initial waves at $\theta = 30^\circ$, and 9% for $\theta = 60^\circ$. These changes in the ion temperature anisotropies are related to the preferential heating of the alpha particles and the perpendicular cooling for the protons in the decreasing fluctuations of the transverse magnetic field. The relative drift speed slowly decreases within the simulation time of 800 proton gyroperiods. It decreases with approximately 8% for the case of parallel waves, 15% for the case of oblique waves propagating at $\theta = 30^\circ$, and is significantly reduced by 34% for the case of highly oblique wave propagation at $\theta = 60^\circ$. We should note that all cases above should be stable from the point of view of linear Vlasov instability. Figure 8 illustrates the temporal evolution of the parallel (along) and the perpendicular (across the ambient magnetic field) components of the ion temperature for protons (top row) and α particles (bottom row) throughout the simulation time. The figure is complementary to Figure 7 and shows the preferential ion heating for alpha particles in the electromagnetic field of the nonresonant initial wave spectra. The plots show that the α particles acquire high perpendicular temperatures for all cases with an estimated growth rate of

65%–72% as we transition from the parallel to the highly oblique initial wave spectra. The parallel temperatures for the alpha particles are also significantly enhanced, and their increment throughout the simulations varies from approximately 50% for the initial parallel waves through 70% for the slightly oblique case and reaches as high as 120% for the case of highly oblique initial wave spectra at $\theta = 60^\circ$. The dominant parallel heating in the latter case leads to parallel temperature anisotropies for the α particles with $T_{\perp}/T_{\parallel} < 1$. The intense alpha particle heating in both parallel and perpendicular directions can be attributed to nonresonant particle scattering and absorption of the initial low-frequency transverse fluctuations, as well as to Landau and cyclotron damping of resonant α -cyclotron waves generated via nonlinear cascade in the course of evolution. The most prominent parallel heating of the alpha particles in the case of highly oblique waves is due to the parallel electric field associated with the electromagnetic Alfvén-cyclotron waves at oblique wave propagation. In the strictly parallel case only longitudinal waves can carry electric field fluctuations, and the initial spectra of transverse Alfvén-cyclotron waves cannot contribute to the electric field and density fluctuations, unless nonlinear wave-wave interactions, such as parametric instabilities or more complicated wave couplings, take place. Ion heating by parametrically unstable Alfvén-cyclotron waves has been discussed, for example, in Araneda et al. (2009) and Maneva et al. (2014). The protons are less affected by the low-frequency initial wave spectra than the

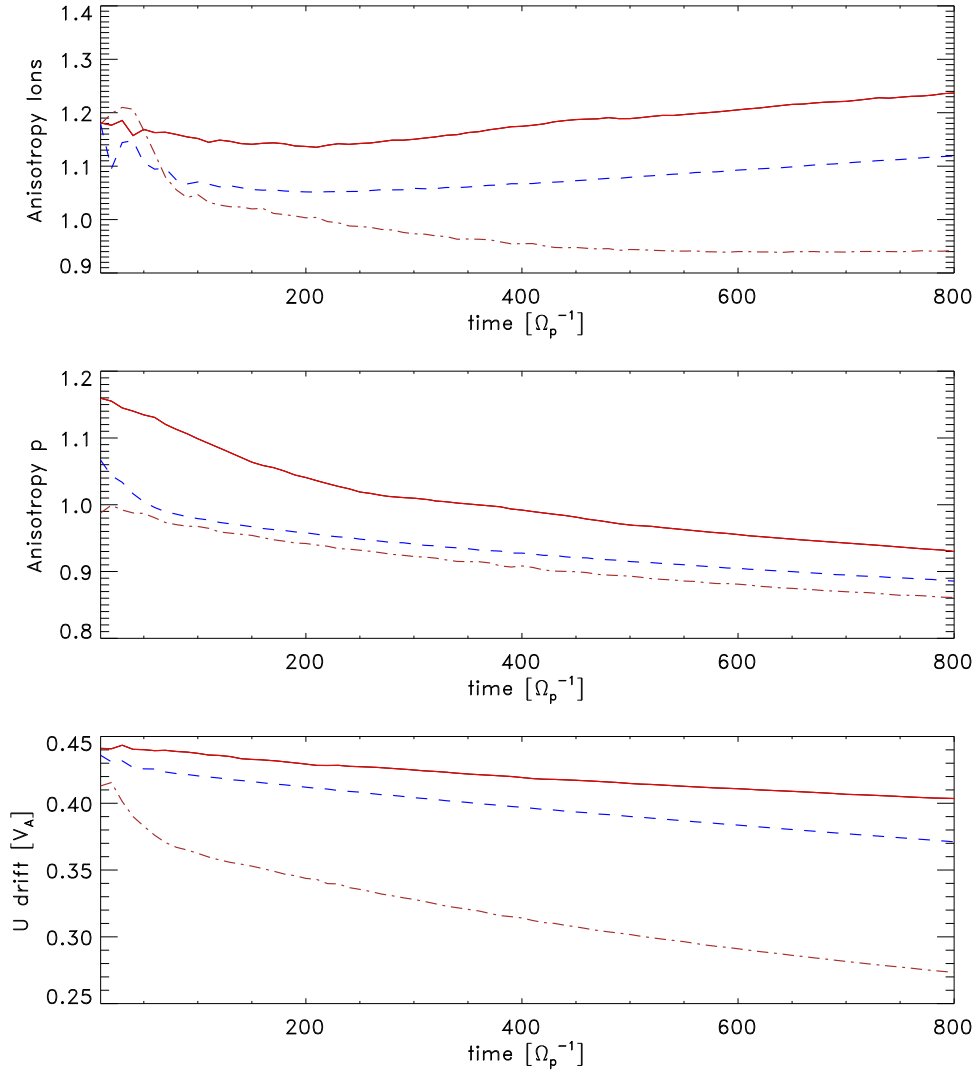


Figure 7. Temporal evolution of the ion temperature anisotropies $T_{\perp i}/T_{\parallel i}$ and relative drift speed for the case of parallel (solid red line) and oblique wave propagation at $\theta = 30^\circ$ (blue dashed line) and $\theta = 60^\circ$ (brown dot-dashed line).

α particles. Their parallel temperature throughout the simulation time increases within 5%–8%, and their perpendicular temperature decreases adiabatically by 10%–14% with the decay of the initial transverse magnetic spectra.

To better estimate the relevance of the individual sources of free energy for the heating of the alpha particles, we need to express the magnetic fluctuations, temperature, and relative drift speed in the same dimensional or dimensionless units and compare their relative change with respect to each other. All simulation quantities presented in the figures have been expressed in dimensionless units, but can be easily converted to dimensional ones. The dimensional magnetic energy is trivially related to the dimensionless one by $\delta B^2/2\mu_0 = (\delta B/B_0)^2 B_0^2/(2\mu_0)$. The thermal energy for the two ion species can be expressed as $1/2(\rho_\alpha V_\alpha^2)$ and $1/2(\rho_p V_p^2)$, where the mass density is $\rho_i = n_i m_i$. The normalized temperatures plotted in Figure 8 correspond to $(m_i/m_p) V_{th,i}^2/V_A^2$, with $V_{th,i}$ being the thermal speed of the i th ion species. Taking into account our definition of the Alfvén speed in terms of the total electron density and the proton mass, as well as the selected number density for the alpha particles, the dimensional temperatures can be expressed as $\tilde{T}_\alpha = 0.03(m_\alpha/m_p)(V_{th,\alpha}^2/V_A^2)(B_0^2/(2\mu_0))$ and $\tilde{T}_p =$

$0.94(V_{th,p}^2/V_A^2)(B_0^2/(2\mu_0))$. From these estimates it becomes clear that 1% change in magnetic energy corresponds to approximately the same change in thermal energy for the protons, while it releases much more energy for the heating of the alpha particles (1% change of the magnetic energy could lead to approximately 33% increase in the thermal energy of the alpha particles). To compare the changes of the relative drift speed to the change in ion thermal energy, we can use several approximations. We could introduce an average mass density $\rho = 1/2(\rho_\alpha + \rho_p)$ and compute $1/2(\rho V_{\alpha p}^2)$, or we can compute $1/2(\rho_i V_{\alpha p}^2)$ for each species. Owing to the low number density of the alpha particles, the average mass is mainly carried by the protons (even though the bulk velocity for the alpha particles is much higher than the bulk velocity of the protons), and then 30% change of the relative drift speed would lead to approximately 9% change in the magnetic energy and the proton thermal energy. Alternatively, we can introduce the reduced mass, $M = m_p m_\alpha / (m_p + m_\alpha)$, and use it to calculate the relative drift speed change separately for the two ion species. If we follow this approach and calculate $1/2(n_i M V_{\alpha p}^2)$ for the two ion species, then the observed 30% variation in the relative drift speed could be approximately compared to 6%–7% change in the proton thermal energy and a

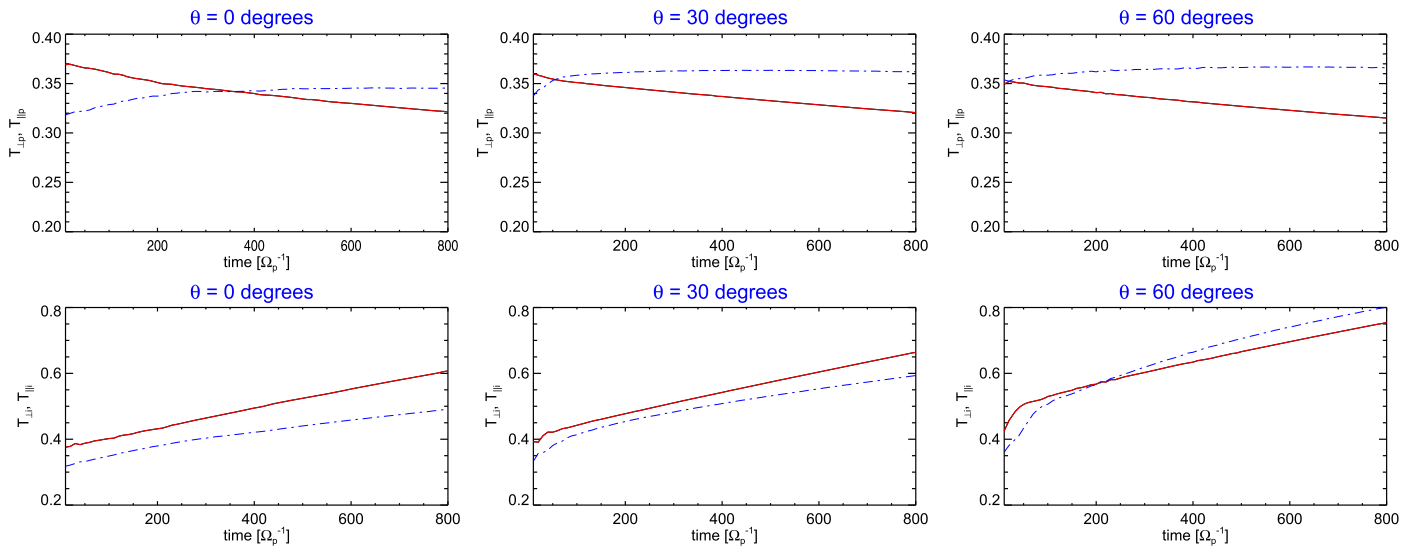


Figure 8. Evolution of the parallel (dot-dashed blue lines) and perpendicular (solid red lines) ion temperature components for protons (top row) and α particles (bottom row). The different panels show the different perpendicular heating for the alpha particles induced by the three different initial wave spectra of parallel (first column) and oblique waves at $\theta = 30^\circ$ (second column) and at $\theta = 60^\circ$ (third column).

similar change in the thermal energy for the alpha particles. Considering that the total magnetic field throughout the simulations decreases with more than 60% for all angles of propagation and the alpha particles are heated by more than 100% in the highly oblique case, the analysis above concludes that the strong preferential heating of the alpha particles in our simulations must originate primarily from wave-particle interactions and absorption of magnetic energy.

Figures 9 and 10 describe the final-stage (at $\Omega_p t = 800$) ion velocity distribution functions for protons and α particles for the three different angles of propagation of the initial Alfvén-cyclotron wave spectra. The left panel shows the case of initially strictly parallel waves, the middle panel describes oblique propagation at $\theta = 30^\circ$, and the right panel shows the evolved velocity distributions for the highly oblique case with $\theta = 60^\circ$. The bottom rows show the particle counts along the parallel direction (solid lines), overplotted with the best bi-Maxwellian fit (dotted lines). The center of all velocity distribution functions is shifted as a result of the presence of relative drift between the α particles and protons. The figures indicate that the final distributions for both ion species deviate from the drifting bi-Maxwellian approximation. Owing to the perpendicular cooling and the parallel heating, the velocity distributions for the proton species are elongated at all propagation angles with a distinct onset of a beam formation in the parallel case. The scattering at the front side (the right-hand side) of the distribution becomes less prominent in the case of highly oblique wave propagation, and at $\theta = 60^\circ$ the proton distributions can be sufficiently well described within the drifting bi-Maxwellian approximation. The velocity distribution functions for the α particles show strong signatures of particle scattering at the front part of the distributions for all propagation angles and at the tail of the distribution functions in the cases of oblique wave propagation at $\theta = 30^\circ$ and $\theta = 60^\circ$. This scattering is caused by the wave-particle interactions with the turbulent spectra and relates to the observed perpendicular and parallel heating for the alpha particles. The distributions are skewed, and throughout the simulation time we observe prominent forward beam formation

for both protons and the α particles in the case of initial parallel wave spectra and a combination of forward and inverse α beam formations (referring to the positive/front and negative/back side of the velocity distribution function) in the case of oblique wave propagation at $\theta = 30^\circ$ and $\theta = 60^\circ$. In the highly oblique case the backscattering by the oblique waves dominates and a dominant inverse α beam is formed. Similar beam formations through particle scattering by initial parallel wave spectra and monochromatic waves have been reported earlier in previous hybrid simulation studies (e.g., Maneva et al. 2013, 2014, 2015). Figure 11 shows the parallel component of the normalized magnetic field at the final stage of the simulation, $\Omega_p t = 800$, for the three different initial propagation angles. The scale is given in units of the magnitude of the constant magnetic background, whose contribution has not been subtracted from the fluctuations. The top panel shows the generated magnetic field fluctuations in the parallel direction for the case of initially parallel wave propagation with $\theta = 0^\circ$. The initial strictly parallel transverse wave spectra include no variations along the direction of the background field and the observed small levels of fluctuations at the final stage are entirely due to the generation of oblique waves. In the case of oblique wave propagation at $\theta = 30^\circ$ there is a structured redistribution of the initial magnetic field fluctuations in real space, enhancing the wave power in some regions of the box, for instance, at $x > 300d_i$ and $y < 100d_i$ (see Figure 5). In the case of initial wave propagation at $\theta = 60^\circ$ the final-stage magnetic field also shows a structured form different from the shape of the initially constructed fluctuations. In both cases in the course of evolution the waves start to form more coherent patterns. The total amplitude of magnetic field fluctuations significantly decreases over time for all three cases. A plausible explanation for the change of the initial wavefront orientation is the energy transfer between parallel and oblique modes. Thus, the highly oblique initial wave spectrum evolves transferring energy to the parallel modes, and the quasi-parallel spectrum evolves transferring energy from quasi-parallel to more perpendicular modes. Figure 12 depicts the evolution of the initial magnetic field power in the Fourier space for the case of

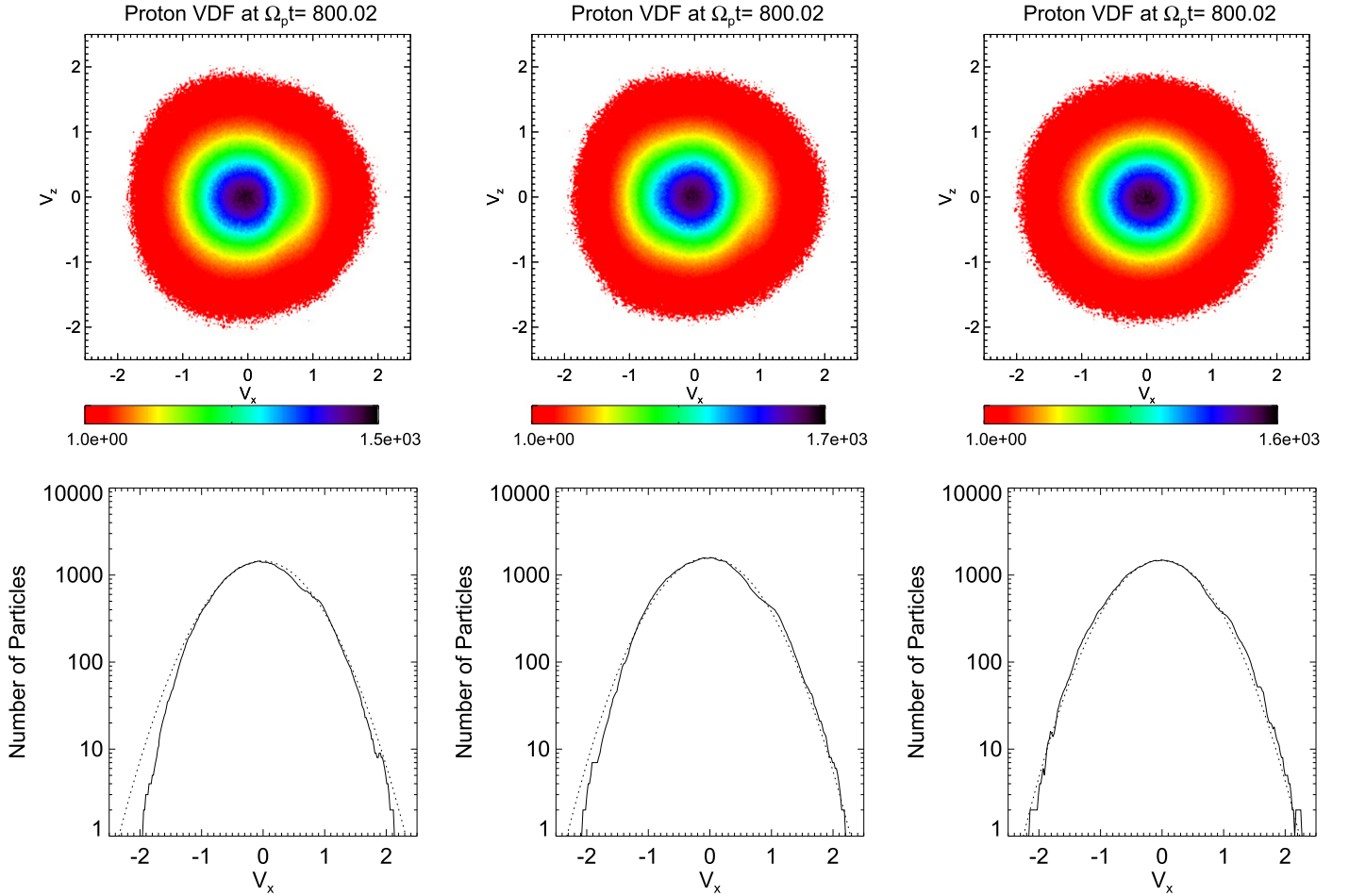


Figure 9. Snapshots with the velocity distribution functions for protons for initially parallel waves (left panel) and oblique waves at $\theta = 30^\circ$ (middle panel) and at $\theta = 60^\circ$ (right panel).

oblique initial wave spectra propagating at $\theta = 30^\circ$. The figure depicts the strong depletion of the initial wave power with significant direct cascade in perpendicular direction particularly visible around $k_y = k_\perp = 0.36 \Omega_p / V_A$ and $k_y = k_\perp = 0.46 \Omega_p / V_A$. We should note that the initial power is transferred to a wide range of perpendicular wavenumbers, which are not initially present in the system (note that the initial range of perpendicular wavenumbers at $\theta = 30^\circ$ is $k_\perp \in [0.13, 0.26] \Omega_p / V_A$, as shown on Figure 4). In addition, a substantial fraction of the wave power has cascaded toward strictly perpendicular wavenumbers, which implies the generation of dispersionless wave modes propagating at $\theta = 90^\circ$. As visible in the next figures, this process persists for the case of initial spectra at $\theta = 60^\circ$, where the wave energy is transferred to even smaller transverse scales with higher perpendicular wavenumbers. Apart from the perpendicular cascade, the spectra also show both direct and inverse cascade in the parallel direction. The direct cascade happens at an earlier stage before the end of the simulations and generates resonant modes, which are quickly absorbed by the alpha particles. Figures 13–15 provide information about the nonlinear evolution of the initial wave spectra and the dispersion of the turbulence-generated waves in the course of the simulations. The simulation dispersion relation is computed from the power of the Fourier-transformed magnetic field fluctuations in space and time. To compute the dispersion in the parallel direction,

the Fourier transformation is taken at a fixed position, from the vertical domain of the initial wave spectra. The same procedure is used in the computation of the wave dispersion in the perpendicular direction. To compute the dispersion, we have used the temporal information throughout the end of the simulations. The magnetic field power along the vertical line perpendicular to k_\perp in the bottom panel of Figure 13 is due to the strong parallel waves, plotted in the top panel. The horizontal lines with $\omega < 0.5 \Omega_p$ may relate to ion Bernstein modes for the α particles, while higher-frequency proton-Bernstein modes are highly damped. The figure indicates the generation of strong forward-propagating fast modes, as well as strong backward-propagating Alfvén-cyclotron modes at a wavenumber range spreading beyond the wavenumbers of the initial parallel wave spectrum. Low-frequency backward-propagating Alfvén waves are undamped solutions of the parallel linear Vlasov dispersion relation when sub-Alfvénic drifts are present. Here they are excited through nonlinear wave-wave interactions in the presence of the initial wave spectra and the induced ion temperature anisotropies. The generation of these background-propagating waves continues at oblique wave propagation, but they are limited within lower wavenumber range, and at $\theta = 60^\circ$ their signature is overcome by strong backward-propagating fast modes, as visible from the next figures. The forward-propagating fast modes are less prominent for the case of $\theta = 30^\circ$ and are overcome by

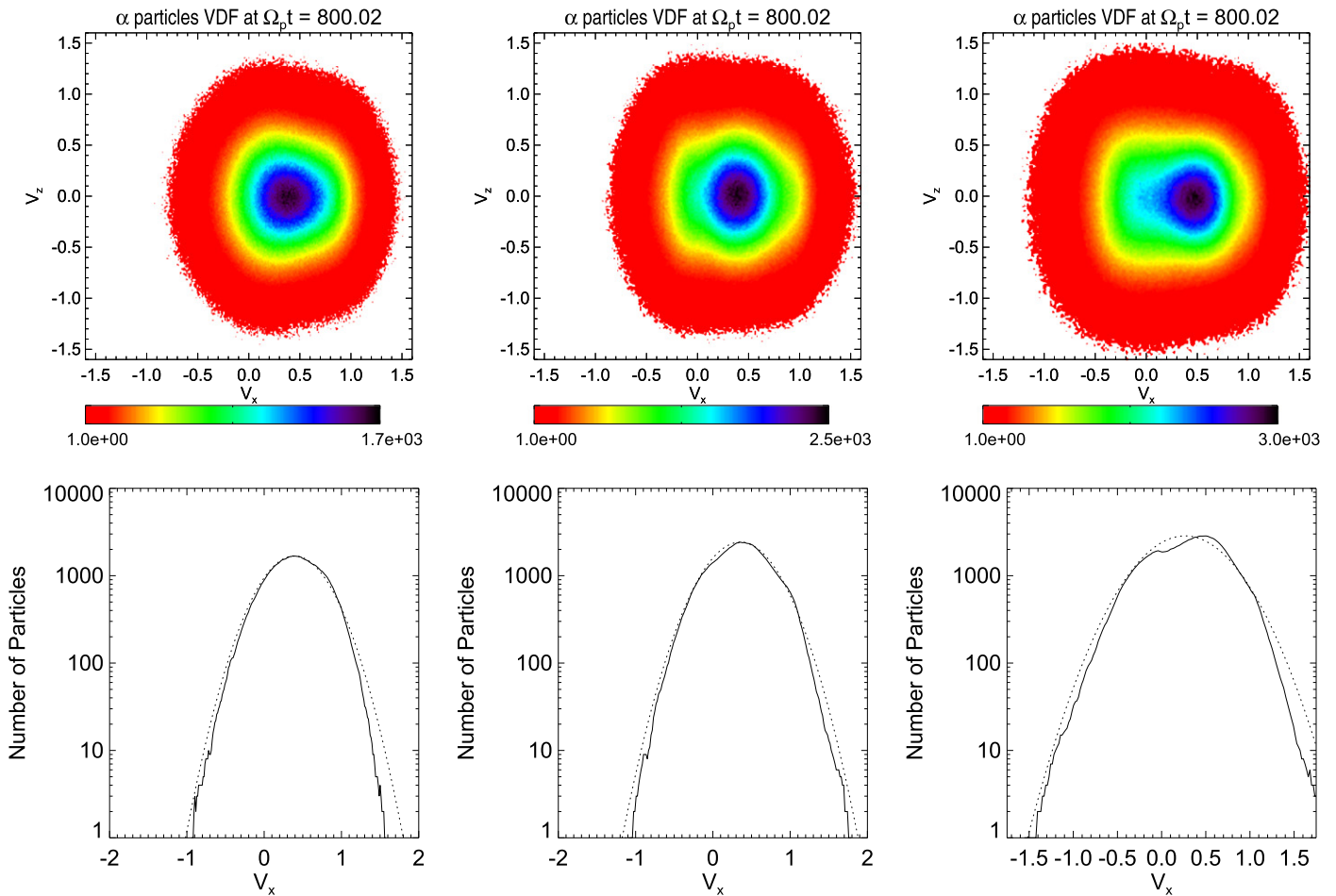


Figure 10. Snapshots with the velocity distribution functions for α particles for initially parallel waves (left panel) and oblique waves at $\theta = 30^\circ$ (middle panel) and at $\theta = 60^\circ$ (right panel).

backward-propagating fast modes when $\theta = 60^\circ$. The zero-frequency lines at the bottom panels in Figures 14 and 15 describe nonpropagating entropy waves, generated by the density and pressure fluctuations. In an isotropic and slightly anisotropic plasma with sub-Alfvénic drifts within the specific wavenumber range, linear instability theory predicts that these waves should be stable or fully damped. The presence of the relative drift speed reduces the damping of the slow modes, and the entropy waves and the pressure fluctuations associated with the initial wave spectra further enhance the entropy waves' wave power. Thus, the positive growth rate of the entropy waves is a nonlinear effect, which can be attributed to the presence of the initial Alfvén-cyclotron wave spectra. Looking at the complementary spectral plot shown in Figure 12 and the explanations thereafter, the zero-frequency modes can be attributed to the generation of strictly perpendicular Alfvén waves, as well as beating of counter propagating oblique Alfvén waves.

6. DISCUSSION AND CONCLUSIONS

Parallel and obliquely propagating ICWs are ubiquitous in the solar wind plasma and have been observed at various heliocentric distances. Furthermore, differentially streaming nonthermal ions with beams in their velocity distribution functions and different temperatures along and across the direction of the ambient magnetic field are common features of

the fast solar wind streams. Yet, the correlations between the existing plasma waves and the particle properties in the multi-species solar wind are still poorly understood, and further theoretical and simulation studies directly related to observations are needed. In this paper we present the results from an observation-driven 2.5D hybrid simulation study, as a first attempt to distinguish the relative role of parallel versus oblique wave propagation for the generation and evolution of ion temperature anisotropies and relative drifts in the fast solar wind streams. The initial state is based on *Wind* data from 2008 January 16, 15:19:09 UT in the undisturbed wind at 1 AU, and we study the evolution of the expanding solar wind plasma parcel, which cannot be followed by the spacecraft. We have investigated the evolution of a stable region within the velocity distribution functions of proton and α particles in anisotropy- β space. The initial isotropic distributions of the alpha particles interact with the wave spectra and acquire slightly higher perpendicular temperatures. In contrast, the protons cool down, as there is no available wave energy at the proton scales. Our findings show that the alpha particles are preferentially heated by the magnetic fluctuations for all initial wave spectra considered here, where protons are prone to perpendicular cooling, as expected by the double-adiabatic expansion model predictions. At the final stage of the simulations the strength of the magnetic fluctuations decreases as the magnetic power is converted into thermal energy for the alpha particles. The main assumption in this work is the type of initial wave spectra

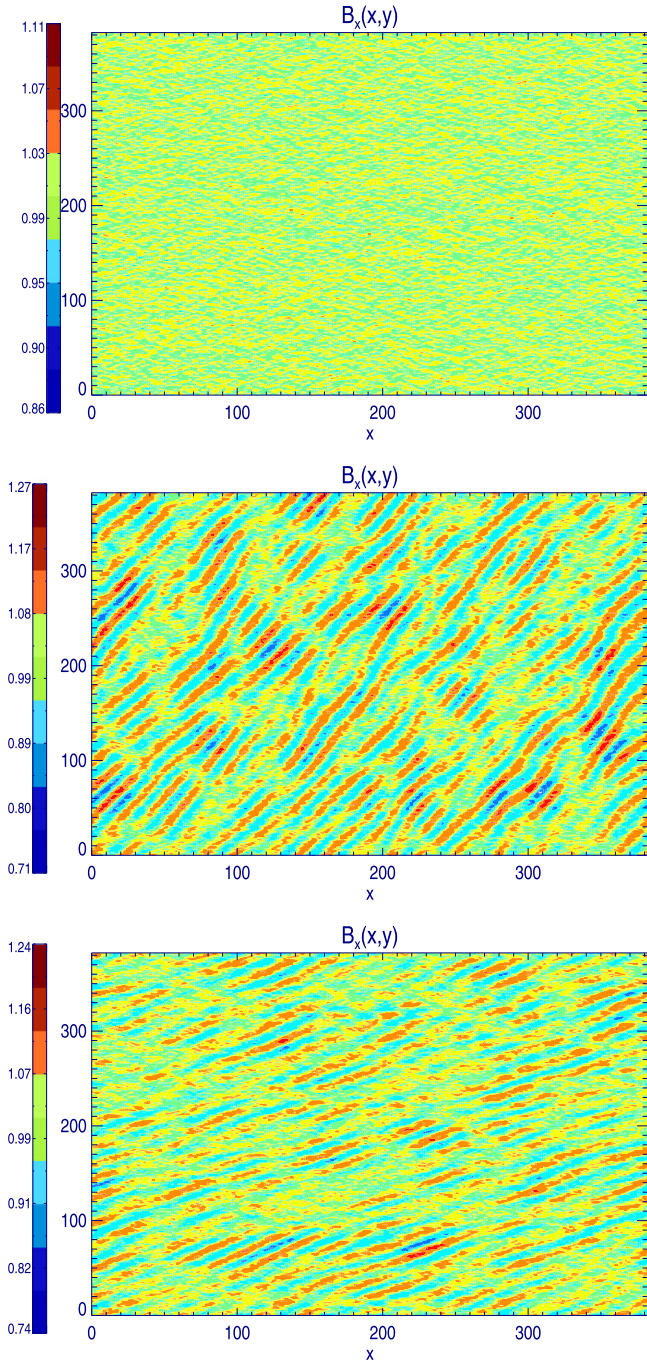


Figure 11. 2D magnetic field fluctuations along the background field at $\Omega_p t = 800$ as a function of normalized coordinates in configuration space. The top panel relates to oblique fluctuation formed from the initially parallel wave spectra. The middle panel shows the end-state magnetic field variations for the case of oblique waves propagating at $\theta = 30^\circ$. The bottom panel represents the reconfiguration of the magnetic field fluctuations for the case of initial oblique wave propagation at $\theta = 60^\circ$.

considered here. Since we are interested in the properties of the alpha particles and try to couple the physics at fluid and ion scales, for the initial broadband wave spectra we have chosen low-frequency left-hand circularly polarized Alfvén-cyclotron waves. We exclude the effect of the high-frequency (much higher than the proton-cyclotron frequency) whistler and KAWs, which interact mainly with high-speed electrons and proton beams. Instead, our study is focused on the direct

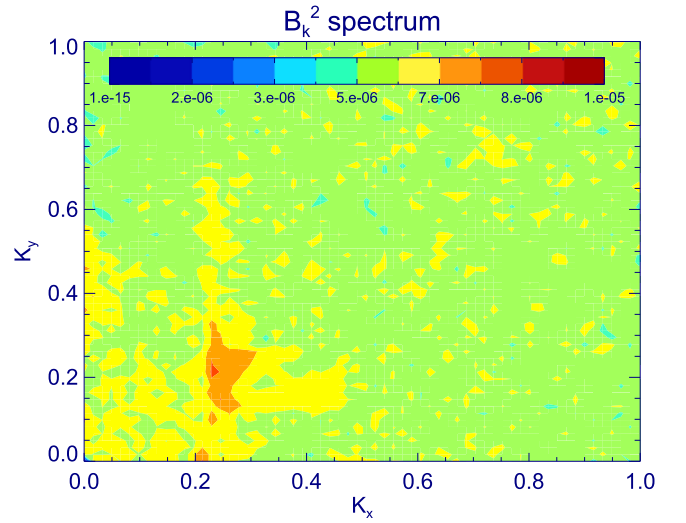


Figure 12. Power spectra of the evolved magnetic field fluctuations at the final stage of the simulations at $\Omega_p t = 800$. The figure shows the square of the magnetic field fluctuations as a function of normalized parallel and perpendicular wavenumbers for the case of initially obliquely propagating waves at $\theta = 30^\circ$.

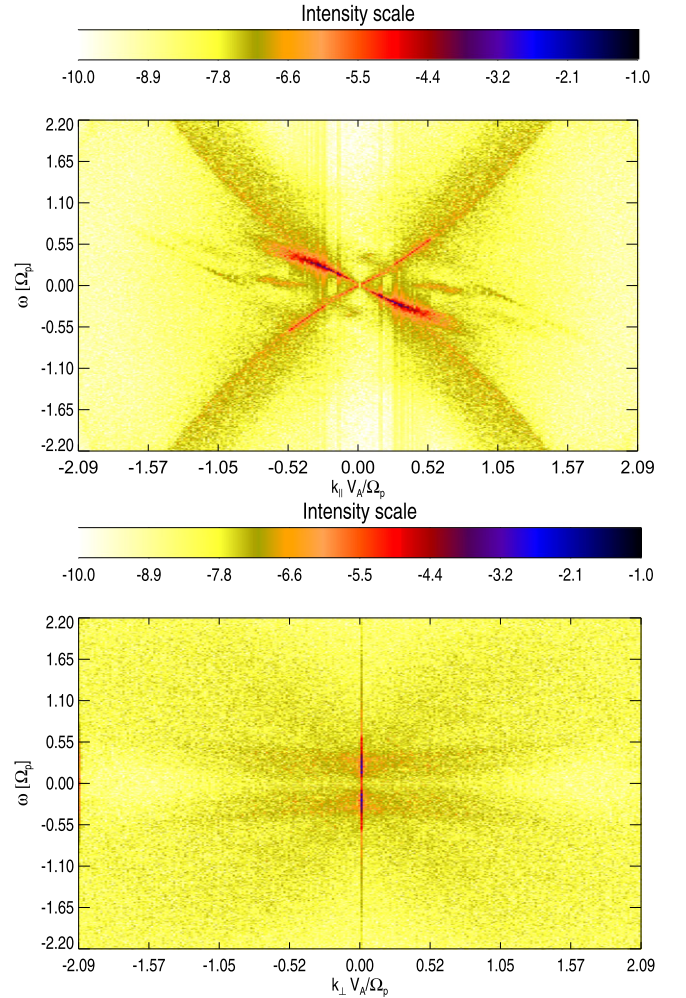


Figure 13. Power spectra of the magnetic field fluctuations in Fourier space at the final stage of the simulations at $\Omega_p t = 800$. The figure shows the square of the magnetic field fluctuations in logarithmic scale as a function of normalized frequency, parallel (top panel) and perpendicular (bottom panel) wavenumbers for the case of initially parallel wave propagation.

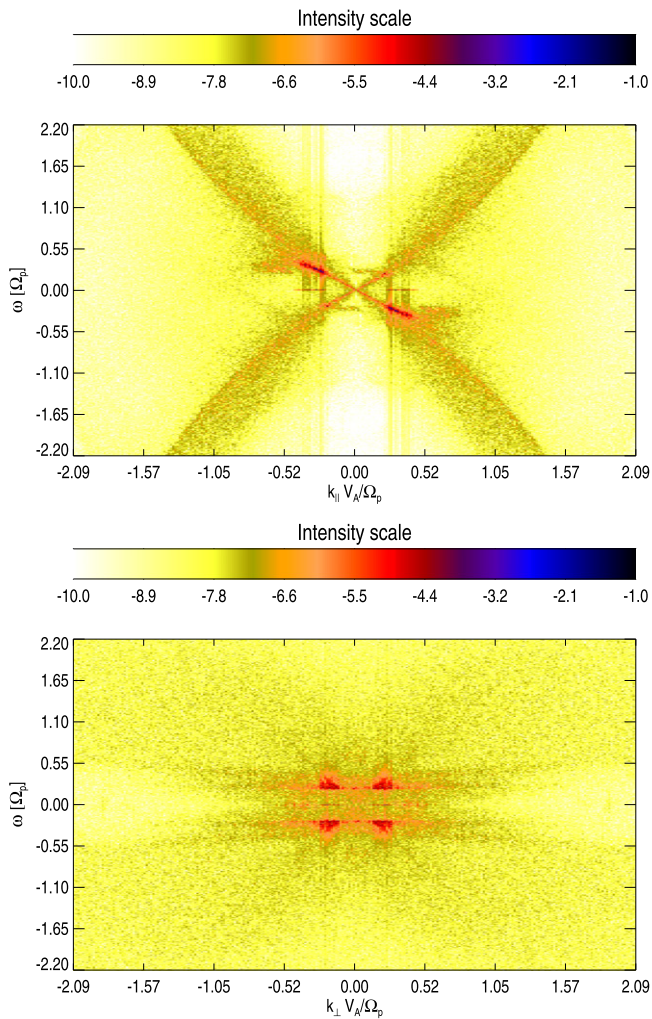


Figure 14. Power spectra of the magnetic field fluctuations similar to Figure 13, but for the case of oblique waves propagating at $\theta = 30^\circ$.

cascade of wave power from the intermediate fluid regime down to the kinetic ion scales and neglects the effects of inverse cascade at scales much smaller than the proton skin depth or their possible contributions to the energy transfer to the ion scales. Yet, apart from the direct cascade toward the smaller ion scales, our simulations do show signatures of inverse cascade toward larger fluid scales with the generation of strictly perpendicular waves as clear from the magnetic field power spectra plot at the end of the simulations. In this study we have performed 2.5D hybrid simulations to investigate the heating of alpha particles by broadband spectra of oblique Alfvén-cyclotron waves and the evolution of an initial relative drift speed between the protons and the alpha particles. We have compared the ion heating corresponding to different angles of propagation (0° , 30° , and 60°) and found prominent perpendicular heating for the alpha particles at all propagation angles. For the selected spectral range of the initial wave spectra for the simulations we find that the α particles are most prominently heated in both the parallel and perpendicular direction by the highly oblique Alfvén-cyclotron waves propagating at $\theta = 60^\circ$. The parallel heating for both species increases with the propagation angle, and the parallel temperature of the alpha particles in the highly oblique case prevails over the perpendicular one, so that they acquire

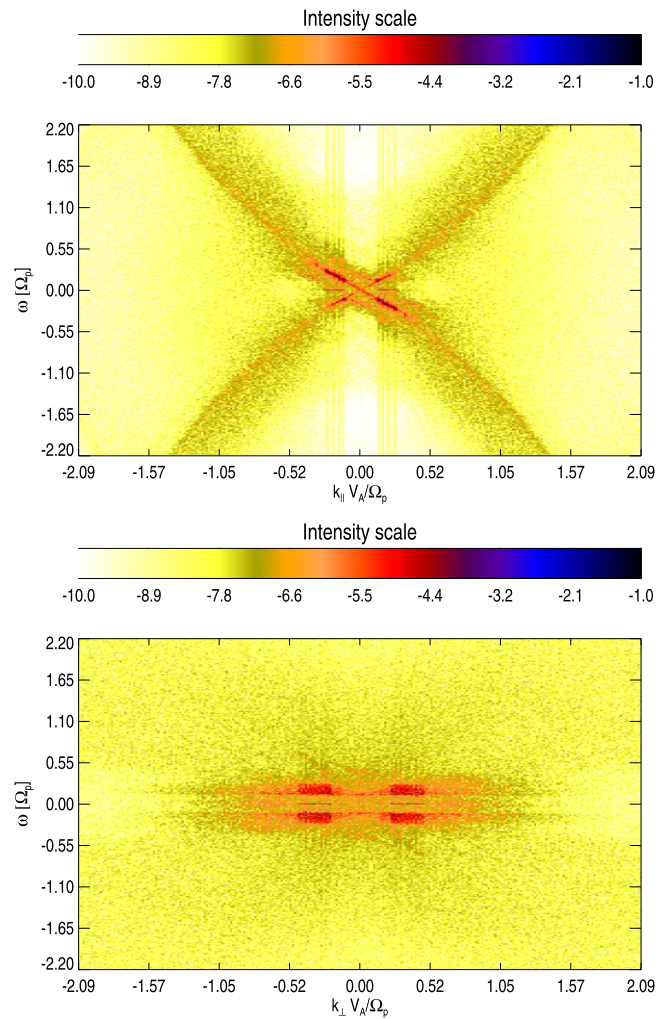


Figure 15. Same spectral plot for the magnetic field fluctuations as in Figures 13 and 14, but for the case of initial oblique wave propagation at $\theta = 60^\circ$.

parallel temperature anisotropy with $T_\perp/T_\parallel < 1$. Since the perpendicular heating of the alpha particles remains at a similar level for all propagation angles, the alpha particles have the highest anisotropy for the case of initially parallel-propagating waves. For all cases considered here the initially isotropic protons experience moderate cooling in the perpendicular direction, and slight heating in the parallel direction, the net effect of which leads to the generation of parallel temperature anisotropies. Contrary to the linear instability theory predictions, the initial relative drift speed $V_{\alpha p} = 0.44 V_A$ decreases in time throughout the simulations. It remains quasi-conserved in the parallel case and gets further reduced in the oblique cases with a magnitude change of over 30% when $\theta = 60^\circ$. As part of this study we have also looked at the effect of a gradual solar wind expansion with $\varepsilon t_0 = 10^{-5}$. As expected by construction, the expansion leads to perpendicular cooling for both species; however, considering the realistically small value of the expansion parameter, it does not have a significant effect on the ion heating within the simulation times considered here. The current study is concentrated on warm plasmas with $\beta_p \approx 0.33$, but this result is expected to hold even stronger when applied to the lower plasma β conditions in coronal holes

and the fast solar wind near the Sun, where $\tilde{\beta} \sim 10^{-2}$. We should note that the $\tilde{\beta}$ used in the simulations directly relates to the standard definition of plasma β once the ion-to-proton mass ratio and ion-to-electron density ratio are taken into account. At the initial state $\beta_p \approx \tilde{\beta}_p$ and $\beta_\alpha \approx 0.12\tilde{\beta}_\alpha$. The initial wave spectra slightly modify the initial plasma $\tilde{\beta}$ and β , introducing up to 11% difference between their parallel and perpendicular values in some cases. The final plasma β values are anisotropic for both species in all cases. The parallel proton $\tilde{\beta}_{\parallel p}$ slightly increases, while the perpendicular $\tilde{\beta}_{\perp p}$ decreases up to 13%. The parallel $\tilde{\beta}_{\parallel \alpha}$ increases by approximately 40% in the case of parallel wave propagation and by more than 100% in the case of highly oblique propagation. The perpendicular $\tilde{\beta}_{\perp \alpha}$ increases by approximately 40% and 90% in the case of parallel and highly oblique wave propagation, respectively.

As a next step we investigate the power spectra of the magnetic fluctuations. We find limited turbulent cascade in the perpendicular direction for the parallel case, accompanied by a strong inverse cascade in parallel direction toward smaller wavenumbers. For the oblique cases we find prominent additional oblique mode generation in a wider range of wavenumbers than the initial values, as well as the generation of strictly perpendicular waves. Right-hand polarized fast waves are being generated in all simulation cases with strong forward-propagating modes in the parallel case and strong backward-propagating modes in the highly oblique case. Strong forward-propagating ion-cyclotron modes are excited in the parallel case, as they are expected normal modes in drifting two-ion-species systems. In the cases of initial oblique wave propagation the initial Alfvén-cyclotron wave spectrum is strongly damped and generation of entropy modes/perpendicularly propagating Alfvén waves is observed. We should note that the total magnetic energy is gradually depleted throughout the simulation time for all three initial angles of wave propagation, and it decreases approximately by 70% from the initial value. Simultaneously, the alpha particles are continuously heated on average by 69% in the perpendicular direction. This analysis suggests that the major part of the alpha particle heating is due to the external wave energy sources and cannot be simply attributed to redistribution of kinetic energy between the protons and the α particles. A further source for the ion heating can be a nonlinear magnetosonic instability induced by the initial drift speed in the presence of the wave spectra. The value of the initial relative drift is within the stable regime for linear plasma instabilities, yet it decreases up to 30% in the highly oblique case. This contributes to the energy source of the initial wave spectra and provides an additional energy source needed for the strong parallel heating with more than a 120% increase of the parallel temperature for the alpha particles. However, as discussed in the results section, a 30% drop in the relative drift speed relates to a 6%–7% increase in the thermal energy of the alpha particles, and the main source for their heating is absorption of the initial magnetic spectra. Another source of free energy is the gradual perpendicular cooling of protons. Note that the relative decrease of the proton temperature anisotropy, which varies between 13% and 20%, is due to a combination of parallel heating and perpendicular cooling. The drop in proton anisotropy is most prominent for the parallel case and weakest for the highly oblique case, whereas the heating of the alpha particles is strongest for the highly oblique case. This is a clear indication that the observed heating of alpha particles is related

to the initial wave spectra, rather than a pure energy exchange between the two ion species.

Looking into the velocity distribution functions, we find strong signatures of wave scattering with prominent forward or inverse alpha beams being formed by $\Omega_p t = 800$. The proton scattering by the oblique waves leads to no clear beam formations. This can support the conclusion that stronger proton beams are formed by trapping in the ion-acoustic wave potential generated in the nonlinear evolution of the initial wave spectra or relate to the fact that it is more difficult to accelerate protons by low-frequency waves in the finite- β plasma considered here in comparison to the lower- β regimes close to the Sun, where the wave amplitudes are also expected to be stronger. We should note that for the limited wavenumber range considered in our study the initial spectral slope of the fluctuations does not play an important role for the nonlinear wave energy transfer and the related ion heating/cooling and acceleration or deceleration. Simulation tests with more gradual initial magnetic field power slope, for instance, $\alpha = -1$, show similar results to the ones presented here. This work does not include post-simulation wave polarization analysis; however, based on previous experiments and the presented wave dispersion analysis in Fourier space, we believe that right-hand-polarized waves are being generated in the course of evolution and the final state consists of a mixture of left- and right-handed waves. In a next improvement of the present analysis we will broaden the spectral range of the initial fluctuations and include different polarization states. As a future perspective we plan to implement a 2D spectrum of a general type, avoiding the rotation of parallel spectra. We would also consider low-frequency right-hand-polarized KAWs and whistlers and investigate the amount of energy they can deposit to the ions before ultimately dissipating at the electron scales.

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