The SERMON project: 48 new field Blazhko stars and an investigation of modulation-period distribution

M. Skarka,[1, 2, 3] J. Liška[2, 3] R. F. Auer[3] Z. Prudil[2] A. Juráňová[2] and Á. Sódor[1]

[1] Konkoly Observatory, MTA Research Centre for Astronomy and Earth Sciences, Konkoly Thege Miklós út 15-17, H–1121 Budapest, Hungary; e-mail: `marek.skarka@csfk.mta.hu`
[2] Department of Theoretical Physics and Astrophysics, Masaryk University, Kotlářská 2, 611 37 Brno, Czech Republic
[3] Variable Star and Exoplanet Section of the Czech Astronomical Society, Valašské Meziříčí, Vsetínská 78, Valašské Meziříčí, 757 01, Czech Republic

ABSTRACT

Aims. We investigated 1234 fundamental mode RR Lyrae stars observed by the All Sky Automated Survey (ASAS) to identify the Blazhko (BL) effect. A sample of 1547 BL stars from the literature was collected to compare the modulation-period distribution with stars newly identified in our sample.
Methods. A classical frequency spectra analysis was performed using PERIOD04 software. Data points from each star from the ASAS database were analysed individually to avoid confusion with artificial peaks and aliases. Statistical methods were used in the investigation of the modulation-period distribution.
Results. Altogether we identified 87 BL stars (48 new detections), 7 candidate stars, and 22 stars showing long-term period variations. The distribution of modulation periods of newly identified BL stars corresponds well to the distribution of modulation periods of stars located in the Galactic field, Galactic bulge, Large Magellanic Cloud, and globular cluster M5 collected from the literature. As a very important by-product of this comparison, we found that pulsation periods of BL stars follow Gaussian distribution with the mean period of 0.54 ± 0.07 d, while the modulation periods show log-normal distribution with centre at $\log(P_\mathrm{m} \, [\mathrm{d}]) = 1.78 \pm 0.30$ dex. This means that 99.7 % of all known modulated stars have BL periods between 7.6 and 478 days. We discuss the identification of long modulation periods and show, that a significant percentage of stars showing long-term period variations could be classified as BL stars.

Key words. Stars: variables: RR Lyrae – Stars: horizontal-branch – Methods: statistical

1. Introduction

Together with ultra-precise space measurements from the *Kepler* and *CoRoT* telescopes, the large-scale surveys play crucial role in investigation of long-term amplitude and frequency (phase) modulation in RR Lyrae stars known as the Blazhko (BL) effect (Blažko 1907). Accurate space data helped uncover various dynamical processes present only in modulated stars (for example, period doubling (Szabó et al. 2010; Kolláth, Molnár & Szabó 2011) and additional radial and non-radial modes (Benkő et al. 2014; Szabó et al. 2014)). Data from the ground-based surveys, on the other hand, are very useful for studying large samples statistically and investigating general and common characteristics of modulated stars (e.g. Alcock et al. 2003; Soszyński et al. 2011).

Both approaches are important in their own way because the nature of the BL effect has not been fully explained yet. The true percentage of modulated stars is still unknown (see e.g. Jurcsik et al. 2009; Kovács 2015). Even more interestingly, it is not clear why some stars show modulation and some with very similar physical properties do not (e.g. Skarka 2014b). Knowledge about characteristics of the BL effect is still insufficient to give a reliable answer to the question of if and how the pulsation period is connected with the length and amplitude of the modulation. Jurcsik, Sódor & Váradi (2005) showed that

the modulation amplitude of short-pulsation-period RR Lyrae stars can be higher than for long-period stars. Benkő et al. (2014) and Benkő & Szabó (2015) suggest a monotonic dependence between the strength of amplitude modulation and the period of the BL cycle (the longer the cycle, the larger the amplitude). While the upper limit of the modulation length is still unclear, the bottom limit is probably defined by the upper limit of the rotational velocity of a star. RR Lyrae stars with short pulsation periods can have short modulation periods, while long-pulsation-period stars can only have a BL cycle that is longer than about 20 d (Jurcsik et al. 2005).

Several studies report on the mode switching and (or) disappearing (reappearing) of the BL effect, suggesting that the phenomenon could be of temporal nature (see e.g. Szeidl 1976; Sódor et al. 2007; Goranskij, Clement & Thompson 2010; Jurcsik et al. 2012; Plachy et al. 2014). Many stars also show variations of their modulation cycles (see e.g. Guggenberger et al. 2012; Benkő et al. 2014; Szabó et al. 2014). In some BL stars, for example in RR Lyrae itself, the modulation reappears cyclically during an additional four-year cycle (Detre & Szeidl 1973; Le Borgne et al. 2014; Poretti et al. 2016). There are many more interesting features of the BL effect. Extensive reviews of the BL effect are provided in recent overviews by Szabó (2014), Kovács (2015), and Smolec (2016).

In his study based on the brightest RRab stars observed by the *All Sky Automated Survey* (ASAS) (e.g. Pojmański 1997, 2001) and *the Super Wide Angle Search for Planets* (Super-WASP; Pollacco et al. 2006), Skarka (2014a)revealed the short-comings of current automated procedures dedicated for the search for the BL effect and the need for permanent supervision of all steps when analysing noisy ground-based data dominated by various aliases. His study also showed that there are many undiscovered modulated stars even in previously studied data sets. Because frequency analysis and searching for the manifestations of the BL effect in many targets individually is a very time-consuming task, we establish a new group involving professionals, students, and amateur astronomers that aim to analyse data from various surveys. The project was called SERMON which is the abbreviation of *SEarch for Rr lyraes with MOdulatioN*.

In this paper we focus on RRab stars observed by ASAS that are brighter than 13.5 mag in maximum light to investigate their frequency spectra trying to reveal the BL effect. Because the ASAS data have special characteristics and need special handling, we describe these data in detail in Sect. 2. We also describe the sample selection in this section. Analysis, criteria for the detection of the BL effect and discussion about the sample stars are described in Sect. 3. In Sect. 4 we investigate the distribution of modulation periods. The results are discussed in Sect. 5, the content of the paper is summarized in Sect. 6.

2. The ASAS data

The ASAS-3 survey is performed by small-aperture telescopes that are primarily dedicated to search for new variables among stars with declination below +28° (technical details can be found in Pojmański 2001). The telescopes regularly scan the sky, and typically observe a particular object once per a few days with a three-minute exposure in *V* filter.

The publicly available data typically contain several hundreds of data points, which are marked with flags from A (the best quality) to D (the worst quality). The photometry often consists of several blocks in the data files because each star is usually observed in a few different fields. Magnitudes with errors are provided in five apertures with diameter between 2 and 6 pixels. The time span of the data is about 3300 days in most cases.

The philosophy of the survey has a deep impact on data characteristics. Low sampling rate causes the frequency spectra to be strongly dominated by daily aliases (Fig. 1). Because the angular resolution of the ASAS instruments is very low (about 15 arcsec/pixel), even with the smallest aperture the light of the target star could be influenced by neighbouring star(s), which is (are) closer than 30 arcsec. This fact could seriously bias the data. We discuss this problem in detail in Appendix A. An extensive overview of the ASAS data characteristics can be found in Pigulski (2014).

2.1. The sample selection and data cleaning

This paper represents the extension of the study by Skarka (2014a) with fainter stars and stars with more scattered data. To some extend it complements the analysis by Szczygieł & Fabrycky (2007) who performed an automatic search on the full sample of all types of RR Lyrae stars in the ASAS Catalogue of Variable Stars (ACVS).

The selection process started with the the International Variable Star indeX catalogue (VSX; Watson, Henden & Price 2006,
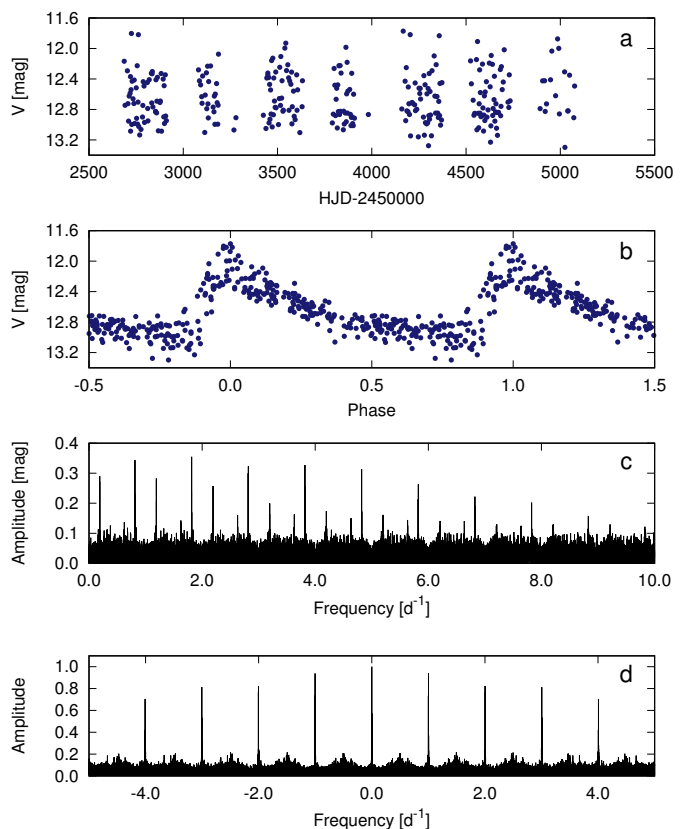
Fig. 1. Typical ASAS data demonstrated on V1124 Her. Panels 'a' and 'b' show the data distribution and data phased with the basic pulsation period. Panel 'c' shows corresponding frequency spectrum; panel 'd' shows the spectral window.

version from 27th October, 2014), in which 1760 stars of RRab type with maximum light brighter than 13.5 mag and declination below +28° were listed. The declination limit is given by the observational constraints of the ASAS survey. The magnitude limit was chosen according to several tests as the most limiting for reliable analysis because the scatter steeply increases with decreasing brightness. The corresponding photometric error of a data point at 13.5 mag is about 0.15 mag (Pojmański 2002).

Stars analysed by Skarka (2014a) were ignored. Then we checked for the availability of the data via the online VSX tools. This resulted in 1435 stars that were observed by ASAS. After downloading the data, we obtained a basic idea about the quality of the data, checked for possible blends in the area with diameter of 30 arcsec, and made a note (col. Blend in Tables 1, B.4, B.5, and col. Rem in Table 2). We used only data with flags A and B in aperture with the smallest average photometry error. Possible shifts in the mean magnitude between blocks were ignored because of the poor data sampling and problems with blends. All points that differed more than 1.5 mag from the median value were automatically removed. No other automatic procedure was applied for improving the data to avoid possible misinterpretation because, for example, scatter caused by modulation could be very easily confused with observational scatter.

In the next step, all stars with less than 150 points were discarded. Then we performed a visual inspection of all light curves again, and searched for the basic pulsation period in stars where the value given in VSX was not reliable. In stars where suitable period was found (or was available in VSX) obvious out-

liers were manually deleted. Stars with extremely scattered data were removed from the sample. This kind of pre-analysis gave us an idea about some peculiarities in the objects and allowed us to identify duplicates (Sect. 3.4) and stars of different variable types (Sect. 3.5). Together with these peculiar objects the sample for further analysis contained 1234 stars.

3. Analysis and identification of the BL stars

A mathematical description of simultaneous amplitude and phase (frequency) modulation suggests the rise of an additional peak with modulation frequency (f_m) and equidistant peaks near the basic pulsation frequency components ($kf_0 \pm f_m$; for details see Benkő, Szabó & Paparó 2011; Szeidl et al. 2012). This equidistant spacing is equivalent to f_m. Because the presence of the side peaks in the frequency spectra is the most distinct and convincing evidence of the BL effect, which is also detectable in scattered data, we searched for this feature using PERIOD04 software (Lenz & Breger 2005). Owing to the characteristics of the data, it was necessary to pay special attention to properly identify correct peaks and avoid possible confusion with aliases.

It is likely that variable contamination with parasite light from nearby stars entering the aperture could resemble the amplitude modulation in some of the sample stars. In such cases, the mean magnitude and total amplitude of the light variations varies in time and produces conspicuous peak in the low-frequency range of the frequency spectrum. A visual check of the data itself would not help much because of large scatter and poor sampling, and statistics fails too for the same reasons. Therefore, the only way to reveal this issue is to identify the false modulation peak. No such peak was observed in any of the sample stars, however, we cannot exclude that some of the stars could be influenced by this problem. This is one of the reasons why we give the information about blends in the tables.

A star was marked as a BL star when a peak with $S/N > 3.5$ in the interval of $f_0 \pm 0.2$ c/d was identified. It was found by Nagy & Kovács (2006) that in stars with frequency separation of less than $1.5/TS$ (periods longer than $2/3TS$), where TS is the time span of the data, it is impossible to distinguish unambiguously between BL modulation and long-term continuous period change (instability) of the main pulsation period (i.e. period-change stars, PC class, as described in the next paragraph). Therefore, only peaks with separation larger than $1.5/TS$ were considered a consequence of the BL effect. An example of the frequency spectrum for a BL star is shown in the left-hand panel of Fig. 2, where an equidistant triplet is seen.
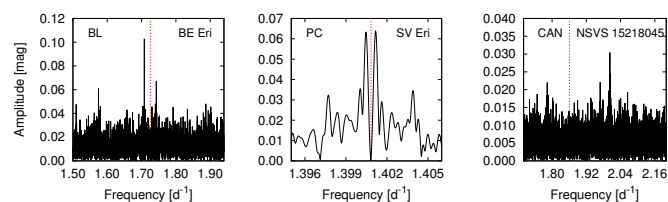
Fig. 2. Typical frequency spectra after prewhitening with the basic pulsation components for BL stars, PC, and candidate stars (from left to the right). The position of the basic pulsation frequency is shown with the red dashed line.

The second category that we adopted is a PC group. Stars assigned to this group are usually those showing either unresolved peaks (peaks with frequency distance lower than $1/TS$),

or it is impossible to prewhiten the peaks in a few consecutive steps (Alcock et al. 2000; Nagy & Kovács 2006). Following our period limit for BL stars, each star with frequency separation lower than $1.5/TS$ was assigned to PC class as well. However, it is sometimes difficult to properly decide about BL/PC category, thus, we adopted these conservative criteria to make sure that PC stars do not contaminate significantly our BL sample (see also the discussion in Sect. 5). A typical frequency spectrum of a PC star is shown in the middle panel of Fig. 2.

Finally, we use candidate category for stars that would belong to the BL group, but the peaks have amplitudes under the detection limit or very close to integer daily frequencies (1 c/d, 2 c/d etc., the right-hand panel of Fig. 2).

3.1. Stars with the BL effect

An online version of the BlaSGalF database[1] containing Galactic field RR Lyrae stars with the BL effect reported in literature contains 58 stars from our sample (version from 29th February, 2016, Skarka 2013). We detected the side peak in only 39 of these stars (see Table B.4). One of these 58 stars belongs to our candidate category, one was found to be of different variability type, and one star is a PC star. In remaining 16 stars we did not detect any peak possibly related to modulation (see Table B.1).

Although we cannot be sure in 100 percent of the stars, we assume, that the nondetection of the BL effect is a consequence of the data characteristics rather then the result of absence of the BL effect. We found two reasons for that. First, the noise level of the residuals after removing significant frequencies is typically between 0.01 and 0.02 mag, which means that sometimes side peaks with amplitude below ~ 0.07 mag do not fulfil $S/N > 3.5$. Second, some of the stars surely showed large-amplitude BL effect at the times when ASAS gathered the data. For example, Jurcsik et al. (2009) easily detected the modulation in the frequency spectra of AQ Lyr and UZ Vir using high-quality photometric observations gathered with a 60-cm telescope between 2004 and 2009. They also did not detect any clear sign of modulation in ASAS and NSVS data that is confirmed by our analysis. In any case, the nondetection of the BL effect in ASAS data should not be considered as a proof of the absence of the modulation.

In 48 stars the BL effect was detected for the first time (see Table 1). In addition to the information about pulsation and modulation periods resulting from our analysis (cols. 2 and 3) and corresponding errors in the last digits in parentheses, we give information about nearby stars closer than 30 arcsec (18 objects with '+' in col. 5 named Blend), and about the number of data points and the time span of the data (cols. 6 N and 7 TS). Four stars denoted with '+' in the last column Rem, in which we detected additional, either unresolved, or unique peaks close to kf_0, are suspected of multiple modulation (category νM from Alcock et al. (2003) and MC from Nagy & Kovács (2006)), or additional long-period variation. Instead of amplitudes of the side peaks, which could be influenced by nearby stars, we give S/N in the fourth column. Errors of periods (in all tables) are 2σ formal uncertainties from the least-squares fitting method.

All data of newly identified modulated stars phased with the modulation periods from Table 1 are shown in Fig. 3. In most of the stars the modulation envelope is well apparent. This figure also shows a variety of envelope shapes and modulation amplitudes. Some of the stars, for example ASAS J181510-3545.7 and V0420 Vel, have very small total amplitudes compared to the

---

Table 1. Stars identified as modulated for the first time. Columns P_{Puls} and P_m give pulsation and modulation periods and corresponding 2σ errors in the last digits; column S/N gives the signal-to-noise ratio; column Blend gives information about nearby stars; columns N and TS give the number of data points and the time span of the data. Stars denoted with '+' in column Rem are suspected of multiple modulation or additional long-period variation.

ID	P_{Puls} [d]	P_m [d]	S/N	Blend	N	TS [d]	Rem
ASAS J103622-3722.1	0.468036(2)	86.2(4)	4.8		553	3289	
CH Aps	0.509710(7)	730(40)	3.9		616	3176	
ASAS J211737+0011.8	0.489135(4)	43.59(14)	4.0	+	295	3235	
OW Aqr	0.655187(2)	172(2)	4.3		404	3254	
V0356 Aqr	0.554594(2)	41.65(6)	4.7		440	3219	
V0522 Cen	0.623662(2)	2100(240)	3.9		582	3187	
V0570 Cen	0.604507(8)	65.35(32)	3.6		470	3170	
BW CMa	0.526804(2)	45.89(12)	6.7	+	530	3299	
ASAS J112830-2429.7	0.678271(2)	11.37(2)	3.9	+	840	3172	
NSV 1888	0.458739(2)	980(60)	4.2		708	3248	
RT Equ	0.444761(10)	106.8(8)	4.2		366	2921	+
BE Eri	0.579536(4)	57.38(12)	6.5		380	3214	
WZ Gru	0.536766(4)	27.35(4)	4.7		360	3261	
ASAS J175143+1708.4	0.435601(4)	6.67(2)	3.8	+	610	2404	
V1332 Her	0.522353(4)	21.59(4)	3.9		435	2404	
ASAS J023020-5908.1	0.572857(4)	29.85(4)	5.3		714	3279	
UW Hor	0.662156(14)	300(8)	5.2		560	3297	
ASAS J094115-1348.8	0.467569(8)	101.0(6)	4.7	+	295	3141	
ASAS J104924-2812.5	0.574553(4)	93.3(6)	5.1		571	3299	
ASAS J110733-2944.1	0.526282(4)	63.48(20)	5.1		592	3154	
ASAS J145702-2642.6	0.586273(8)	75.26(36)	4.2	+	419	3170	
NSV 14546	0.349662(2)	1070(60)	4.4		733	3294	
ASAS J151317-1445.9	0.613006(4)	43.33(8)	6.8	+	588	2965	
CU Lib	0.506956(6)	121.6(8)	4.6	+	473	3175	+
AC Men	0.553409(4)	120.2(8)	4.5	+	380	3288	
ASAS J055109-7219.4	0.558809(4)	168.3(16)	4.4		634	3296	
ASAS J095835-8311.7	0.543469(2)	243(2)	8.4		1344	3294	
BN Oct	0.508267(2)	23.93(4)	5.2		580	3275	
EE Oct	0.444320(2)	175.9(16)	4.5		750	3293	
XZ Oct	0.473867(2)	26.00(4)	4.9		739	3251	
CSS J165728.3+055825	0.592147(2)	406(8)	5.9		343	2959	
FU Pav	0.454554(6)	206(4)	4.0		280	3215	
GV Pav	0.575562(4)	109.4(6)	5.2		481	3215	
NSVS 11726192	0.576882(8)	31.61(6)	4.1		388	3239	
SS Pic	0.494099(2)	61.86(14)	5.9		621	3287	
RR Pyx	0.491078(4)	30.17(4)	5.2	+	878	3295	
AK Scl	0.494987(2)	94.2(6)	5.5		558	3283	
CN Scl	0.585871(4)	1270(80)	4.2		600	3286	+
V0484 Ser	0.511390(4)	119(1)	3.6		324	2946	
ASAS J181510-3545.7	0.579472(4)	127.7(6)	5.2	+	1200	3174	
ASAS J193321-2517.6	0.507410(2)	55.7(2)	4.4	+	438	3132	+
V1069 Sgr	0.478929(6)	206(2)	5.3	+	501	3067	
V2281 Sgr	0.479455(6)	180(2)	3.8	+	360	3271	
ASAS J185332-5133.3	0.611209(8)	432(10)	3.9		466	3155	
ASAS J162246-6723.5	0.543464(2)	85.7(4)	4.7	+	909	3181	
UX Tuc	0.509089(8)	50.25(16)	4.2	+	288	3254	
V0420 Vel	0.605077(2)	66.38(10)	10.2	+	598	3299	
ASAS J073358-6508.1	0.538119(6)	113.2(8)	3.9	+	510	3299	

other stars, which is most likely a result of blends with nearby stars.

### 3.2. Stars with unresolved peaks

Our 22 PC stars are listed in Table 2. The modulation periods suggested by the frequency spacing range from 1900 d in BK Eri to 4000 d in SV Eri. The latter is a star with the most rapid period change known among RR Lyrae stars so far (Poretti et al. 2016). When the data are phased with the suggested modulation periods, DI Aps, QZ Cen, V0584 Cen, NSVS 14009, V0898 Mon, and V0687 Pup show apparent amplitude changes. It may be possible that in a data set with long enough time span, these stars will turn out to be classical BL stars with very long modulation periods (see also Sect. 5 and Fig. 7.)

Short pulsation periods of ASASJ 054843-1627.0, J160612-4319.2 and their low amplitude and phase Fourier coefficients (transformed to I-band filter) suggest that they are of RRc type when compared to characteristics of Galactic bulge RR Lyrae parameters provided by Soszyński et al. (2011).

3.3. Candidates for BL stars

In seven stars we noticed peaks within $f_0 \pm 0.2$ c/d, but we mark them as candidates from the reasons discussed previously. In YZ Aps, NSVS 15218045 (see Fig. 2), ASAS J054238-2557.3, and V1211 Sgr, the suspicious peaks were very close to integer multiples of 1 c/d, which is always problematic. In ASAS J081549-0732.9 and ASAS J194745-4539.7 only peaks with $S/N < 3.5$ were detected.
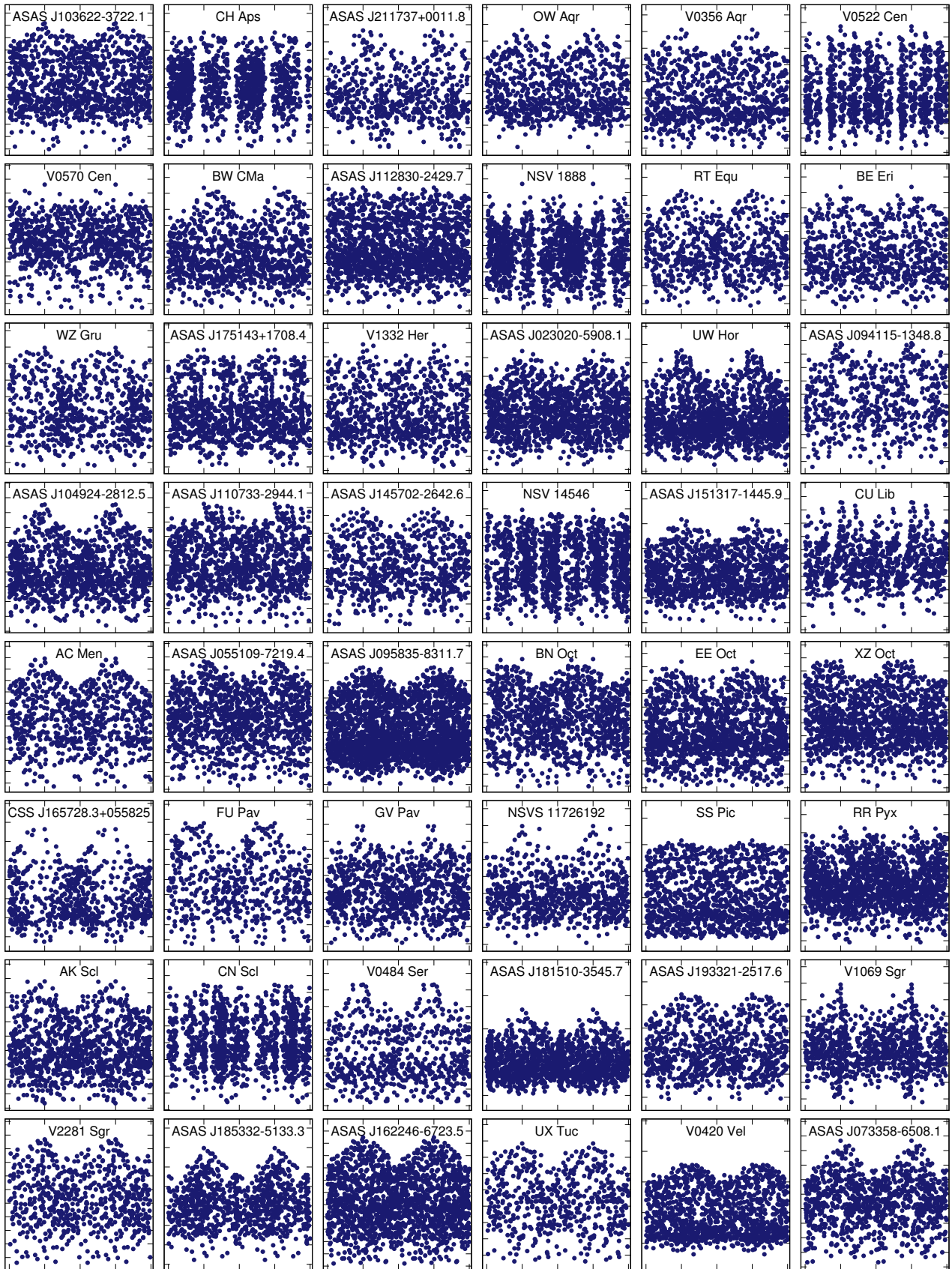
Fig. 3. BL stars phased with modulation periods from Table 1. As the zero epoch we simply took HJD of the brightest data point. Ticks at the ordinate show interval of 0.2 mag, at abscissa the interval is 0.5 in BL phase.

Table 2. Stars showing close, unresolved peaks near f_0 that were marked as PC. The notation of the columns is the same as in Table 1. The star with '*' is known to the literature as a BL star.

ID	P_{Puls} [d]	P_m [d]	S/N	N	*TS* [d]	Rem
DI Aps	0.519187(6)	2700(300)	7.7	509	3212	
ASAS J204705-0919.2	0.508809(6)	2500(200)	6.7	369	3146	+
ASAS J083815-6025.9	0.465751(5)	2400(400)	5.9	550	3293	b
QZ Cen	0.502052(6)	2200(200)	6.1	634	3188	b
V0584 Cen	0.459032(2)	3200(600)	5.0	418	3153	
RS Crv	0.536834(6)	2200(200)	6.0	727	3169	
BK Eri	0.548148(4)	1900(100)	4.9	363	3243	+
SV Eri	0.713861(4)	2800(100)	10.0	536	3296	
NSV 14009	0.563985(6)	3000(200)	6.4	585	3256	
ASAS J054843-1627.0*	0.376709(4)	3000(200)	7.9	544	3299	b
PS Lup	0.471878(2)	3200(400)	5.7	663	3191	+
V0898 Mon	0.541493(12)	2200(200)	5.0	316	3078	b
ASAS J160612-4319.2	0.356944(2)	2800(400)	8.0	639	3194	b
ASAS J165834-8458.6	0.549312(6)	3200(500)	5.1	561	3241	
V0455 Oph	0.453893(2)	2200(400)	3.6	377	2586	b
KV Pav	0.527929(6)	4000(600)	6.9	528	3289	
V0338 Pav	0.530827(6)	2600(400)	4.4	330	3273	b
ASAS J075127-4136.3	0.335716(2)	3000(300)	6.8	610	3296	b
V0687 Pup	0.575194(8)	2500(300)	5.8	483	3246	b
ASAS J200147-2153.8	0.546631(14)	2500(400)	4.3	562	3144	b
ASAS J190713-5205.7	0.526724(2)	3100(600)	4.0	725	3173	
GR Tel	0.611916(4)	2900(400)	4.9	696	3134	b

Notes. Blended stars are marked with 'b' and stars with additional peaks are marked with '+' in the column Rem.

3.4. Duplicates

During our analysis we identified three stars as duplicates because a star with a different ID is located almost at the same coordinates and it has similar pulsation (variability) properties. Actually, this was a trigger to perform a detail analysis of this problem regarding the total VSX catalogue (Liška et al. 2015). Our duplicates are given in Appendix B, Table B.3.

## 3.5. Stars with different variable type

In 24 stars (Table B.2 in Appendix B) we detected no signs of variability, 19 other stars are probably not RRab stars, but we suspect that they are of different variability type (Table B.5, Fig. B.1 in Appendix B). We give suggested variability types that were estimated on the basis of period, light-curve shape, and $B - V$ (Table B.5, Fig. B.1).

3.6. Characteristics of sample BL stars

The percentage of all BL stars in our sample (both known + new detections) is very low, at slightly less than 7 %. The reason is a large scatter that could hider the modulation, rather than a real lack of modulated stars among faint stars. Therefore, this number has nothing to do with the real amount of modulated stars, but only gives information about the detectability of modulation in the data set. Our assumption is reinforced by the look of the magnitude distribution of the sample stars shown in Fig. 4 (see the percentages above bins). The highest number of BL stars was detected between 12.5 and 13.0 mag [2] and decreases steeply in the most populated area with faint stars with scattered data.

In the majority of the new BL stars the modulation is very apparent causing large amplitude changes (e.g. BW CMa, V0484 Ser), while in some of the stars the modulation in amplitude is barely seen (e.g. NSV 14546, SS Pic, see Fig. 3). Even in scat-

[2] The percentage is not the highest in this magnitude range.

tered ASAS data, a wide palette of the shapes of modulation envelopes is apparent. In this sense, one of the most interesting examples is UW Hor, which seems to show non-sinusoidal amplitude modulation. On the other hand, BN Oct and ASAS J162246-6723.5 show nearly sinusoidal variations in amplitude. The BL peak f_m itself was detected only in BL Col, a known BL star. In four new BL stars and three PC stars we detected additional peaks, which could be interpreted as additional modulation, or signs of additional period change. Unfortunately, the low quality of the data does not allow us to investigate these peculiarities in detail.
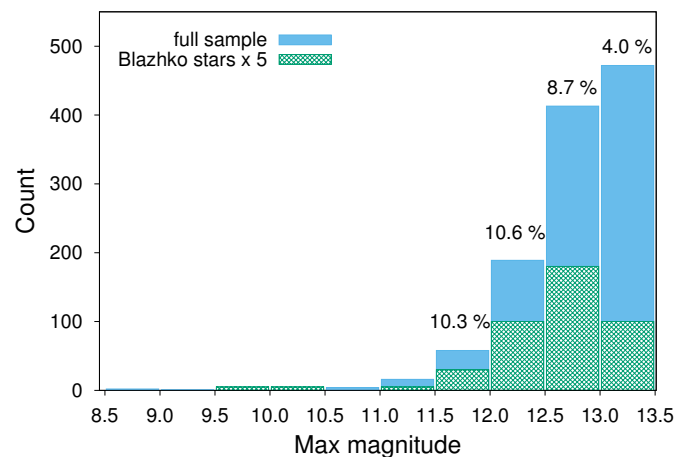
Fig. 4. Magnitude distribution of all sample stars (blue) and BL stars (green grid). The distribution of modulated stars was enlarged five times for better visibility.
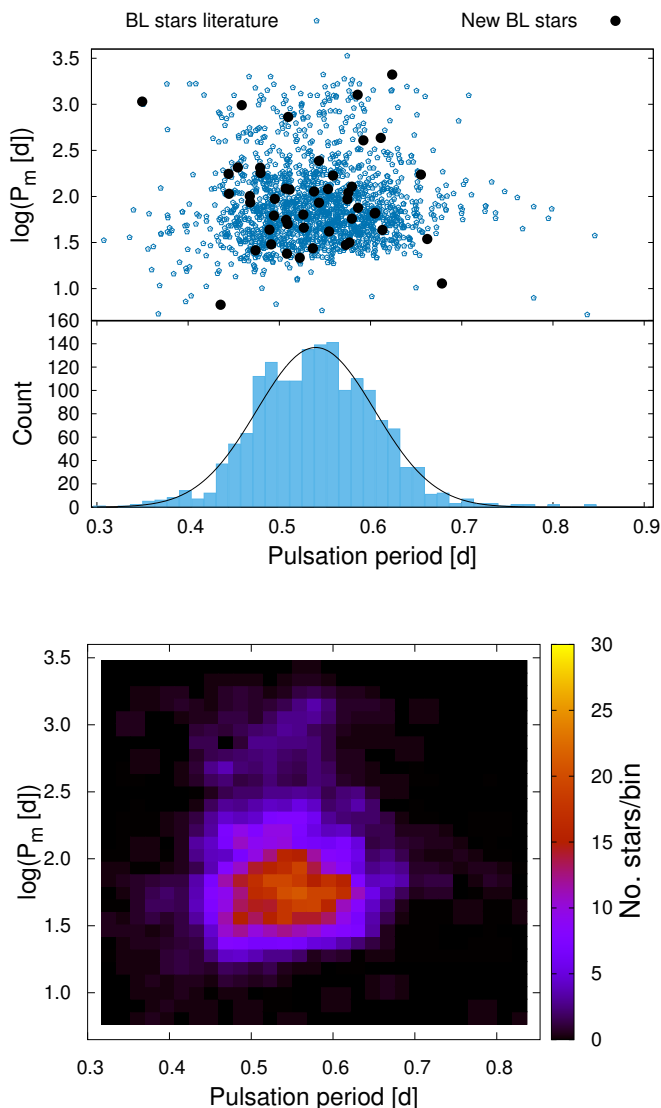
4. Modulation-period distribution

To get an idea about the new BL stars, we collected a sample of 1547 fundamental mode stars with estimated modulation period from the literature. This sample contains RRab stars from the Galactic field (246 stars from the BlaSGalF database; Skarka 2013), the LMC (731 stars; Alcock et al. 2003), the Galactic bulge (19+526 stars; Moskalik & Poretti 2003; Collinge et al. 2006), and globular cluster M5 (25 stars; Jurcsik et al. 2011). We considered all periods in stars with multiple modulation. For example, from RS Boo we have three values on our list. Together with our new detections (48), it amounts 1628 entries, which is, to our knowledge, the most extended sample of modulation periods investigated so far.

The modulation periods of newly identified BL stars (black circles in the top panel of Fig. 5) match the distribution very well, and no sample star has extreme modulation period[3]. Two of our stars with the shortest modulation period, ASAS J175143+1708.4 and ASAS J112830-2429.7, however, seem to significantly differ from the general trend. The BL effect in these two stars should be proved independently. Light curve shape, pulsation period of only about 0.35 d, and its Fourier parameters suggest that NSV 14546 could be an RRc-type rather than RRab-type star.

To test the similarity of our sample with the sample of stars from literature, we performed the two-sample Kolmogorov-

[3] The shortest BL period was detected in ASAS J175143+1708.4 (6.67 d), the longest in V0522 Cen (2100 d), which is exactly the limit for BL stars that we adopted (2/3*TS*).

Fig. 5. BL period as a function of the pulsation period for known and new modulated stars (top), and corresponding density plot (bottom). The bottom part of the top panel shows the distribution of pulsation periods.

Smirnov test. The result is, that the agreement between the two populations cannot be ruled out with probability of 7 %. Thus we can conclude that our sample stars have a similar distribution of modulation periods as BL stars from literature (see the cumulative-distribution functions in the right-hand panel of Fig. 6).

However, the modulation-period distribution in Fig. 5 itself is far more interesting than the distribution of new BL stars[4]. While pulsation periods show a distribution close to the normal Gaussian distribution (the bottom part of the top panel in Fig. 5), modulation periods follow the log-normal distribution. In other words, they show normal distribution when the BL periods are plotted in logarithmic scale, which is nicely seen from the bottom panel of Fig. 5 and histograms in the left-hand panel of Fig. 6. The mean pulsation period of BL stars is 0.54 ± 0.07 d, and the mean modulation period based on the log-normal distri-

---

[4] From here on we do not distinguish between new and known BL stars; all stars create one sample.

bution is 1.78 ± 0.30 dex. This means that 99.7 % (3σ interval) of all BL stars have modulation periods between 7.6 and 478 days.

5. Discussion

The log-normal distribution commonly appears in measured quantities that cannot be negative. Many of natural, social, and technical systems and quantities show such distribution, for example the length of a latent infection (from infection to burst of the disease), species abundance, age of marriage, and probability of failure of mechanical devices (nice examples across the disciplines can be found in Limpert, Stahel & Abbt 2001). The question is why the pulsation periods have Gaussian distribution, while modulation periods show log-normal distribution. Why do the majority of BL stars have modulation periods between 30 and 120 days (1σ interval)? At the present time, it is impossible to answer this question because no correlation between any of the physical characteristics and the length of the modulation period has been found.

The selection effects only have minor impact on the distribution, if any. The majority of used BL stars come from the data sets with time span that is several years long, where the modulation periods of several hundreds of days should be safely detectable. Yearly aliases cause problems with identification only in a very small part of the distribution (see the gap around 350 d in the top insert in the left-hand panel of Fig. 6). Thus we are convinced that the log-normal distribution is real. The very good agreement between the cumulative-distribution function for the periods in this interval and the normal distribution with the same parameters also confirms our assumption (the insert in the right-hand panel of Fig. 6).

It seems that the log-normal distribution only represents the short-period part of the modulation period distribution well (0–400 d), which has a long tail towards longer modulation periods. However, the distribution of the longest BL periods, which is probably much more numerous than our collection shows, is somewhat uncertain from several reasons. First, the entire sample is sharply cut at BL periods of about $\log P_m \sim 3.5$ (~ 3300 d). This limit is given by the extension of the longest data used for modulation period estimation (the ASAS-3 data with maximum time span of about 3300 days). Data from MACHO survey has time span of about 7.5 years. It can be naturally expected that the distribution continues to even longer modulation periods because we do not know the upper limit. What is probably the longest BL effect with period longer than 25 years is reported by Jurcsik & Smitola (2016) in V144 in globular cluster M3.

Recent observational investigations (e.g. Benkő et al. 2014) showed that the BL stars always show both amplitude and phase variations. The BL stars could then be identified through the variations in amplitude even when the modulation period is comparable with *TS*, or longer. For example, V0898 Mon and V0687 Pup from our sample (top panels of Fig. 7) show significant amplitude changes with suggested periods that are a bit longer than $2/3TS$. Although the modulation periods are only very roughly estimated in these stars, the BL-phased light curves suggest, that they could safely be considered BL stars.

Another uncertainty relies in the problems with the BL/PC classification. Although secular period changes and nonstationarity in pulsation period are quite common among RR Lyrae stars (see e.g. Le Borgne et al. 2007), we can expect a significant number of PC stars to be classified as BL stars in long-enough data sets. This is, for example, the case of OGLE-BLG-RRLYR-07697 and OGLE-BLG-RRLYR-06232. These two stars were classified as pure PC stars by Collinge et al. (2006) on the
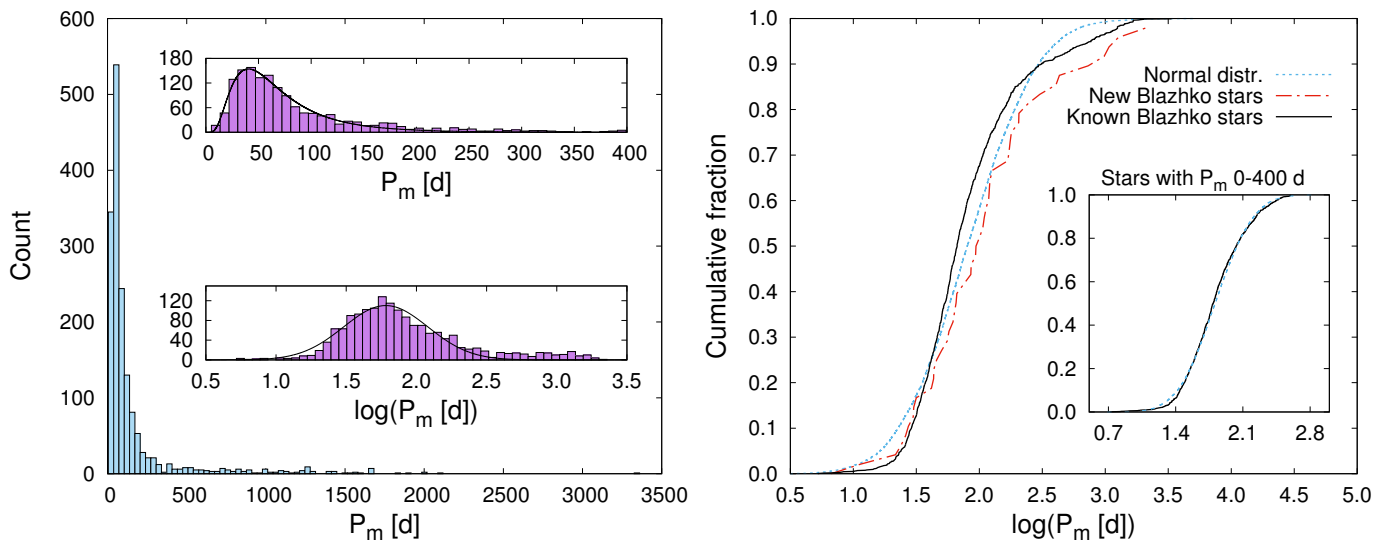
Fig. 6. Distribution of modulation periods (left-hand panel) and cumulative distribution functions (right-hand panel). It is apparent that the overall distribution follows the log-normal distribution. The agreement with the normal distribution for stars with modulation periods between 0 and 400 days shown in the insert in the right-hand panel is very good.
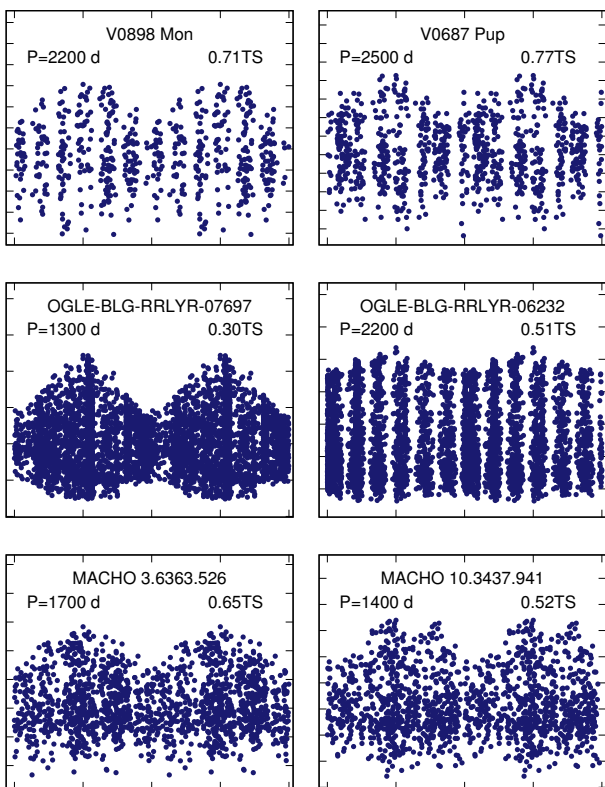
Fig. 7. Stars originally classified as pure PC. There are stars from our sample in the top panel. The middle panels show OGLE-III data demonstrating that in long-enough data sets (original classification comes from OGLE-II data with three-times shorter time span) the stars could be classified as BL stars, not pure PC stars. The bottom panels show two stars that clearly show amplitude modulation with periods shorter than 2/3*TS* suggesting classification as BL+PC, not pure PC. The scales of the axes are the same as in Fig. 3.

basis of the four-year-long OGLE-II data. In OGLE-III data (Soszyński et al. 2011), with a time span of about 12 years, these two stars appeared to be clear BL stars showing significant amplitude modulation with periods of 1300 and 2200 d (the middle panels of Fig. 7). In OGLE-BLG-RRLYR-07697 the side peaks were easily prewhitenable[5]. In OGLE-BLG-RRLYR-06232 we detected additional, unresolved structures very close to f_0, which cannot be fully eliminated, thus, suggesting some additional period variations. According to the criteria adopted in Sect. 3, this star should be classified as both BL and PC star.

Further, we also suppose that stars showing dense patterns close to f_0, which cannot be prewhitened in a few steps, do not necessarily exclude the possibility of being BL stars. In the bottom panel of Fig. 7 we see two stars identified as pure PC stars on the basis of MACHO data (Alcock et al. 2003). Both of these stars clearly show large amplitude modulation with periods shorter than 2/3*TS*. Thus, there is no reason why we do not assign them to the BL class with a possible additional PC component. It would be advisable to strictly distinguish between unresolved peaks close to f_0, which could mean instability of the main pulsation period, and other unresolved peaks near the side peaks, which in turn could mean the nonstationarity of the modulation itself, additional modulation, or some artificially generated periodicity in the data. In such cases the star should be classified as BL star with additional PC, rather then pure PC.

As many previous investigations show (LaCluyzé et al. 2004; Sódor et al. 2007; Detre & Szeidl 1973; Szeidl 1976; Le Borgne et al. 2014; Guggenberger et al. 2012), the BL modulation itself can change on year-long time scale. Although the ASAS and MACHO data sets span rather long time bases, they typically do not contain sufficient number of data points to investigate such details. Therefore, many stars with changing BL modulation might remain undetected or classified as PC stars. These very few examples show that the distribution of modulation periods that are longer than ∼ 1000 d is certainly more populated than we see in Fig. 6.

---

[5] The star shows additional BL modulation with period of 36.48 d.

6. Summary and conclusions

We performed an in-depth study of frequency spectra of 1234 fundamental RR Lyrae stars observed by the ASAS survey, placing emphasis on proper identification of possible modulation. Therefore, we investigated each target individually within a new project SERMON employing professionals, amateurs, and students. We omitted stars with a low number of points (< 150) and scattered light curves with insufficient quality for period analysis. The limiting magnitude was set to 13.5 in maximum light because of large scatter of the data in faint stars.

The criteria for accepting a star as modulated were purely on the basis of the appearance of the frequency spectra. In this sense we establish three groups: *i)* BL stars: stars with peak with $S/N > 3.5$ in the vicinity of f_0, no closer than $1/TS$, but closer than 0.2 c/d; *ii)* BL candidates: stars with one suspicious peak at integer daily frequencies, or with peaks close to f_0 with $S/N < 3.5$; and *iii)* PC: stars with unresolved side peaks or with a separation of the peaks suggesting modulation periods that are longer than $2/3TS$. Altogether we detected 87 BL stars, 48 of these stars for the first time. Even in low-cadence ASAS data with large scatter we observed a variety of possible shapes of modulation envelopes (Fig. 3). We also detected 7 candidates and 22 PC stars. When we put together stars from the ASAS-3 survey analysed by Skarka (2014a) with the sample analysed in this study, the detection efficiency of individual approach is about 12 % in ASAS data, which is more than two times better than in Szczygieł & Fabrycky (2007)[6].

For the comparison of newly identified BL stars, we collected a sample comprising 1547 stars with known BL periods from literature. The distribution of modulation periods of newly identified stars corresponds well to the distribution of periods of known BL stars. As a by-product of this analysis, we noticed that the pulsation periods of BL stars follow Gaussian distribution (mean 0.54 ± 0.07 d), while the modulation periods are distributed log-normally with mean $\log P_m[d] = 1.78 \pm 0.30$ dex. Especially modulation periods in range 0–400 d follow the log-normal distribution well.

Behind its log-normal part ($P_m > 480$ d), the distribution shows an almost constant number of stars per bin, however, this could be simply observational bias caused by the lack of suitable data with long-enough time span. Additional ambiguity comes from the difficulties in properly assigning a star to the PC or BL class. From our discussion about long-term modulation, it follows that stars showing amplitude modulation with period shorter than $2/3TS$ can be safely classified as BL stars, no matter whether there are additional peaks that cannot be prewhitened. These unresolved peaks simply mean that some additional non-stationarity is present in the star, not that the star is unmodulated. Therefore, the classification criteria for PC stars could be revised and summarized in two simple points: *i)* the separation of the side peaks and f_0 is lower than $1.5/TS$, whether or not the peaks are prewhitenable and *ii)* the star shows no apparent amplitude changes. Combined data sets, for example OGLE-III and OGLE-IV, could shed more light on the problem with PC stars and allow us to establish more firm criteria.

The reasons standing behind the log-normal distribution are, unfortunately, unclear. Could there be any connection between the length of modulation period and evolutionary stadium, metallicity, and Oosterhoff groups? These questions would need a deep-in detailed investigation, which is out of scope of our study.

---

[6] It is worth noting that they analysed data with a time span two years shorter than our data sample.

References

Alcock, C., Allsman, R., Alves, D. R., et al. 2000, ApJ, 542, 257
Alcock, C., Alves, D. R., Becker, A., et al. 2003, ApJ, 598, 597
Benkő, J. M., Szabó, R., & Paparó, M. 2011, MNRAS, 417, 974
Benkő, J. M., Plachy, E., Szabó, R., Molnár, L., & Kolláth, Z. 2014, ApJS, 213, 31
Benkő, J. M., & Szabó, R. 2015, European Physical Journal Web of Conferences, 101, 06008
Blažko, S. 1907, Astronomische Nachrichten, 175, 325
Bonnardeau, M., & Hambsch, F.-J. 2015, Information Bulletin on Variable Stars, 6132, 1
Bramich, D. M., Alsubai, K. A., Arellano Ferro, A., et al. 2014, Information Bulletin on Variable Stars, 6106, 1
Chadid, M., Benkő, J. M., Szabó, R., et al. 2010, A&A, 510, A39
Collinge, M. J., Sumi, T., & Fabrycky, D. 2006, ApJ, 651, 197
de Ponthière, P., Hambsch, F.-J., Krajci, T., & Menzies, K., Wils, P. 2013, Journal of the American Association of Variable Star Observers (JAAVSO), 41, 58
de Ponthière, P., Hambsch, F.-J., Menzies, K., & Sabo, R. 2014, Journal of the American Association of Variable Star Observers (JAAVSO), 42, 298
Detre, L., & Szeidl, B. 1973, Information Bulletin on Variable Stars, 764, 1
Goranskij, V., Clement, C. M., & Thompson, M. 2010, in Variable Stars, the Galactic Halo and Galaxy Formation, ed. C. Sterken, N. Samus, & L. Szabados (Moscow: Sternberg Astronomical Institute of Moscow Univ.), 115
Guggenberger, E., Kolenberg, K., Nemec, J. M., et al. 2012, MNRAS, 424, 649
Hajdu, G., Catelan, M., Jurcsik, J., et al. 2015, MNRAS, 449, L113
Jurcsik, J., Sódor, Á, & Váradi, M. 2005a, Information Bulletin on Variable Stars, 5666, 1
Jurcsik, J., Szeidl, B., Nagy, A., & Sódor, Á 2005b, Acta Astron., 55, 303
Jurcsik, J., Sódor, Á., Szeidl, B., et al. 2009, MNRAS, 400, 1006
Jurcsik, J., Szeidl, B., Clement, C., Hurta, Z., & Lovas, M. 2011, MNRAS, 411, 1763
Jurcsik, J., Hajdu, G., Szeidl, B., et al. 2012, MNRAS, 419, 2173
Jurcsik, J., Smitola, P., Hajdu, G., & Nuspl, J. 2014, ApJL, 797, L3
Jurcsik, J., & Smitola, P. 2016, Commmunications of the Konkoly Observatory Hungary, 105, 167
Khruslov, A. V. 2011, Peremennye Zvezdy Prilozhenie, 11,
Kolláth, Z., Molnár, L., & Szabó, R. 2011, MNRAS, 414, 1111
Kocián, R., Liška, J., Skarka, M., et al. 2012, Open European Journal on Variable Stars, 154, 1
Kovács, G. 2005, A&A, 438, 227
Kovács, G. 2015, arXiv:1512.05722
LaCluyzé, A., Smith, H. A., Gill, E.-M., et al. 2004, AJ, 127, 1653
Le Borgne, J. F., Paschke, A., Vandenbroere, J., et al. 2007, A&A, 476, 307
Le Borgne, J. F., Poretti, E., Klotz, A., et al. 2014, MNRAS, 441, 1435
Lenz, P., & Breger, M. 2005, Communications in Asteroseismology, 146, 53
Liška, J., Skarka, M., Auer, R. F., Prudil, Z., & Juráňová, A. 2015, Open European Journal on Variable Stars, 170, 1
Limpert, E., Stahel, W. A., Abbt, M. 2001, BioScience, 51, 5
Moskalik, P., & Poretti, E. 2003, A&A, 398, 213
Nagy, A., & Kovács, G. 2006, A&A, 454, 257
Pigulski, A. 2014, Precision Asteroseismology, 301, 31
Plachy, E., Benkő, J. M., Kolláth, Z., Molnár, L., & Szabó, R. 2014, MNRAS, 445, 2810
Pojmański, G. 1997, AcA, 47, 467

Pojmański, G. 2001, IAU Colloq. 183: Small Telescope Astronomy on Global Scales, 246, 53

Pojmański, G. 2002, AcA, 52, 397

Pollacco, D. L., Skillen, I., Collier Cameron, A., et al. 2006, PASP, 118, 1407

Poretti, E., Le Borgne, J.-F., Klotz, A., Audejean, M., & Hirosawa, K. 2016, Commmunications of the Konkoly Observatory Hungary, 105, 73

Skarka, M. 2013, A&A, 549, A101

Skarka, M. 2014a, A&A, 562, A90

Skarka, M. 2014b, MNRAS, 445, 1584

Smolec, R. 2016, arXiv:1603.01252

Sódor, Á., Szeidl, B., & Jurcsik, J. 2007, A&A, 469, 1033

Sódor, Á., Hajdu, G., Jurcsik, J., et al. 2012, MNRAS, 427, 1517

Soszyński, I., Dziembowski, W. A., Udalski, A., et al. 2011, AcA, 61, 1

Szabó, R. 2014, IAU Symposium, 301, 241

Szabó, R., Benkő, J. M., Paparó, M., et al. 2014, A&A, 570, A100

Szabó, R., Kolláth, Z., Molnár, L., et al. 2010, MNRAS, 409, 1244

Szczygieł, D. M., & Fabrycky, D. C. 2007, MNRAS, 377, 1263

Szeidl, B. 1976, IAU Colloq. 29: Multiple Periodic Variable Stars, 60, 133

Szeidl, B., Hurta, Zs., Jurcsik, J., Clement, C., & Lovas, M. 2011, MNRAS, 411, 1744

Szeidl, B., Jurcsik, J., Sódor, Á., Hajdu, G., & Smitola, P. 2012, MNRAS, 424, 3094

Walker, A. R., & Nemec, J. M. 1996, AJ, 112, 2026

Watson, C. L., Henden, A. A., & Price, A. 2006, 25th Annual Symposium, The Society for Astronomical Sciences, 47W

Wils, P., & Sódor, Á. 2005, Information Bulletin on Variable Stars, 5655, 1

Wils, P., Lloyd, C., & Bernhard, K. 2006, MNRAS, 368, 1757

Zacharias, N., Finch, C. T., Girard, T. M., et al. 2013, AJ, 145, 44

Appendix A: Blend troubles

We demonstrate the problem with blends on CSS J054243.2-114742, where three similarly bright stars and three additional fainter stars are present within area with diameter of 1.5 arcmin = 6 pixels in ASAS (the top panel of Fig. A.1). The data set of this star comprises nine blocks of data points in ASAS-3 database. The middle panel of Fig. A.1 shows how the data from block 9 differs in apertures 1 and 5. It is obvious that in the largest aperture the light contribution of all neighbouring stars is taken into account. This results in higher brightness and lower pulsation amplitude in aperture 5 (6 pixels) than in aperture 1 (2 pixels).
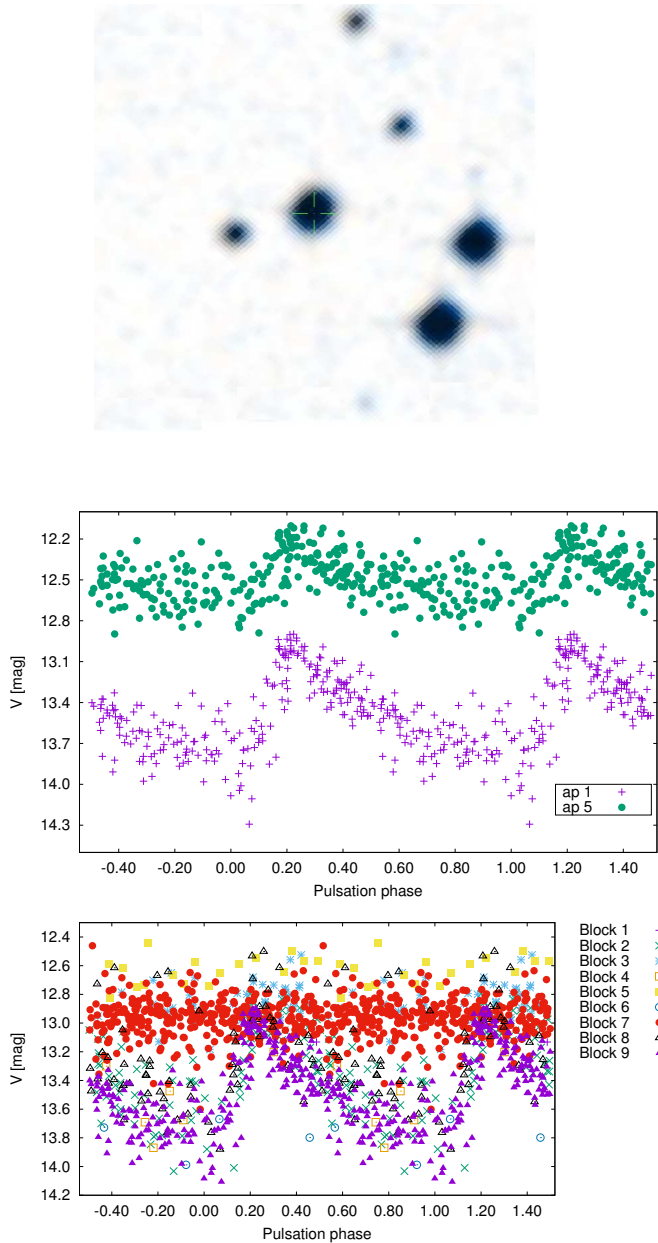
Fig. A.1. ASAS data of CSS J054243.2-114742 (star in the centre of the top left panel) phased with elements given in VSX. The top right panel demonstrates problems caused by large aperture, the bottom panel shows confusing the target with nearby stars in different fields (blocks of data).

An additional problem caused by close stars could arise from the combination of different fields. The bottom panel of Fig. A.1 shows the measurements in aperture 1 for different blocks of the data set. It is apparent that in different fields (blocks 3, 5, 7) the variable star was mistaken for one of the close nonvariable stars. Only blocks 2 and 9 probably show the data for the target star exclusively.

Appendix B: Additional tables and figures

Table B.1. Known BL stars where we did not detect the modulation. The column S/N gives information about S/N of the highest peak detected after removing significant pulsation components in the vicinity of $f_0 \pm 0.2$ c/d.

ID	P_{Puls} [d]	S/N	$P_{\mathrm{literature}}$ [d]	Ref
AL Eri	0.656947(4)	3.3	-	K05
ASAS J030534-3116.1	0.496453(2)	3.2	6.67	SF07
V0461 Her	0.513006(6)	3.0	-	B14
V1162 Her	0.547929(6)	3.3	-	B14
GL Hya	0.505932(4)	3.4	157	B14
V0487 Hya	0.561482(6)	3.2	64.4	B14
ASAS J101200+1921.9	0.482830(10)	3.3	1141.03	SF07
MS Lib	0.441446(8)	3.1	105	B14
AQ Lyr	0.356809(4)	3.2	64.9	J09
V0829 Oph	0.569222(8)	3.1	165	Wa06
V2683 Oph	0.597012(4)	2.9	26.28	SF07
V0606 Peg	0.529658(4)	2.9	26.7	B14
FI Psc	0.531285(6)	3.2	>170	B14
CzeV0245	0.513328(4)	3.0	-	K12
UZ Vir	0.459394(4)	3.3	68.24	J09,S12
ASAS J194502+2434.2	0.845945(22)	3.2	37.56	SF07

Notes. Ref: K05 – Kovács (2005); SF07 – Szczygieł & Fabrycky (2007); B14 – Bramich et al. (2014); J09 – Jurcsik et al. (2009); Wa06 – Watson, Henden & Price (2006); K12 – Kocián et al. (2012); S12 – Sódor et al. (2012).

Table B.2. Stars in which we cannot find any sign of variability or detect any period.

ID	ID
2MASS J06533015+1531072	NSVS 9525696
CSS J163942.1+223202	V0429 Ori
CSS J102417.2+161949	CSS J225023.1+171307
CSS J150035.2-165444	ASAS J004535+0526.1
CSS J150917.1-170252	MM Pup
CSS J153512.2-145339	NSVS 19123091
NSVS 12552884	NSVS 19168136
NSVS 15305366	V0713 Sco
CSS J164808.5+012515	ASAS J185602-1934.6
OGLE BLG-RRLYR-00252	NSVS 19562915
V0771 Oph	OGLE BLG-RRLYR-16496
CSS J053340.7-021532	V0701 Sgr

Table B.3. Stars with duplicate occurrence in VSX (the same coordinates and pulsation properties).

ID	Coord. (J2000)
ASAS J165728+0558.4	16 57 28.42 +05 58 25.5
CSS J165728.3+055825	16 57 28.39 +05 58 25.7
CSS J160717.1-063352	16 07 17.12 −06 33 53.0
V0681 Oph	16 07 15.33 −06 34 18.1
CSS J213502.5+195641	21 35 02.59 +19 56 41.7
ASAS J213502+1956.7	21 35 02.61 +19 56 41.0

Table B.4. Known BL stars from literature detected in our ASAS sample. The notation of the columns is the same as in Table 1.

ID	P_{Puls} [d]	P_{m} [d]	*S/N*	Blend	*N*	*TS* [d]	$P_{\text{literature}}$ [d]	Ref
CK Aps	0.623640(5)	56.4(2)	3.9	+	887	3205	-	K05
KM Aql	0.438191(4)	192(2)	4.7		384	3162	192.2	SF07
PQ Aqr	0.512292(4)	120.3(9)	4.5		362	3144	-	B14
V0354 Aqr	0.529555(6)	182(1)	6.2		398	3291	181.2	SF07
ASAS J172721-5305.9	0.435434(2)	58.5(2)	4.8		615	3171	58.66	SF07
ASAS J123812-4422.5	0.523543(6)	1540(120)	6.6	+	534	3162	1307.7	SF07
ASAS J135813-4215.1	0.523179(3)	145(1)	5.3	+	492	3197	146.01	SF07
GW Cet	0.516652(4)	84.2(4)	4.1		462	3278	84.99	SF07
ASAS J053830-3554.4=BL Col	0.590761(4)	40.53(6)	6.3		583	3289	40.05	K11
AI Crt	0.502901(4)	63.0(2)	5.8		622	3176	63	WS05
V0365 Her	0.613178(10)	40.8(1)	4.5	+	224	2396	40	W06
V1124 Her	0.551018(4)	38.92(6)	6.4	+	296	2388	38.8	B14
RV Hor	0.572496(4)	79.7(3)	5.8	+	512	3264	79.81	SF07
ASAS J093731-1816.2	0.520920(8)	87.7(3)	6.0		320	3294	87.73	SF07
ASAS J141025-2244.8	0.639874(4)	1700(100)	7.3		567	3186	1556.66	SF07
V0430 Hya	0.496828(2)	58.5(3)	4.6		363	3109	56.6	B14
V0486 Hya	0.508644(4)	53.4(2)	4.0		405	3286	18.5	B14
V0543 Hya	0.598263(6)	58.8(2)	5.2		668	3299	59	B14
CZ Ind	0.605147(4)	133(1)	4.1		462	3257	133.38	SF07
V0550 Peg	0.493047(13)	21.98(4)	4.0		219	3235	348.58	SF07
V1820 Ori	0.479040(6)	27.95(2)	6.5		364	2625	27.917	dP13
ASAS J185719-6321.4	0.412008(2)	63.0(2)	4.6		496	3149	61.39	SF07
CS Phe	0.484395(4)	61.2(2)	5.9		530	3287	62.5	WS05
NSV 0539 = DR Dor	0.460421(2)	40.10(4)	7.0		515	3256	40.17	SF07
AL Pic	0.548618(6)	34.13(6)	5.0	+	404	3287	34.07	dP14
NSV 1856	0.516080(2)	742(14)	8.4		1018	3248	786.91	SF07
FR Psc	0.455680(4)	51.47(8)	5.8		305	3240	51.31	dp14
NSV 3331=V0714 Pup	0.494182(2)	119.4(7)	7.2		893	3272	116.96	SF07
BT Sco	0.548718(4)	39.11(8)	4.6		508	3167	78	WS05
AR Ser	0.575213(8)	1325(60)	6.3		463	3167	109	WS05
V5659 Sgr	0.379706(4)	45.58(8)	5.9	+	488	3134	45.7	WS05
ASAS J153830-6906.4	0.622472(6)	117.9(6)	5.9	+	660	3209	1702.13	SF07
EP Tuc	0.615003(6)	63.0(2)	4.0		451	3274	63	WS05
BQ Vir	0.637027(4)	50.8(1)	4.4		805	3160	-	K05
V0354 Vir	0.595045(8)	49.0(2)	4.0		407	3168	59	W06
V0551 Vir	0.446854(8)	53.0(2)	4.6	+	289	2941	48	W06
V0574 Vir	0.474391(8)	25.71(4)	4.1	+	227	3099	26.3	B14
V0585 Vir	0.601611(4)	94.1(5)	4.7	+	574	3183	93.8	B14
V0586 Vir	0.682773(4)	134(1)	4.0	+	517	3174	132	B14

Notes. Ref: K05 – Kovács (2005); SF07 – Szczygieł & Fabrycky (2007); B14 – Bramich et al. (2014); W06 – Wils et al. (2006); K11 – Khruslov (2011); WS05 – Wils & Sodor (2005); dP13 – de Ponthière et al. (2013); dP14 – de Ponthière et al. (2014); BH15 – Bonnardeau & Hambsch (2015).

Table B.5. Stars of different variability type originally classified as RRab stars in VSX. Colour index is a rough value taken from Zacharias et al. (2013).

ID	*P* [d]	Type	Blend	*B−V* [mag]	*N*	*TS* [d]	Remarks
NSVS 14248877	0.99842(4)	E:	+	1.2	344	2939	
ASAS J134527-3616.8	4.77429(10)	CEP		1.1	164	1008	
ASAS J142204-3404.3	0.773789(6)	CEP		1.0	141	998	
ASAS J020345-1729.1	4.089(12)	CEP		1.1	455	3296	
NSVS 15925156	1.06135(4)	CEP		1.1	580	3293	jump in amplitude at 2453500
ASAS J205934+1812.7	3.24610(4)	CEP	+	1.2	491	2422	
ASAS J040615-3050.0	1.32465(4)	CEP		0.9	593	3300	
ASAS J144908-1150.8	119.907(2)	M	+	1.6	494	3174	
CSS J153535.3-155042	0.121131(8)	DSCT		0.9	499	2966	amplitude only 0.1 mag
ASAS J202746-2850.5	0.408462(4)	RRC/EW		0.2	701	3250	additional period with 90 d
ASAS J053552-0508.2	2.82362(4)	CEP		1.2	754	3298	
ASAS J183333-6654.0	1.924273(4)	CEP	+	1.0	1066	3162	
CSS J222816.3+192154	0.140082(14)	DSCT		0.4	292	2366	amplitude only 0.1 mag
ASAS J085816-3022.1	1.90949(4)	ELL	+	1.8	598	3291	amplitude only 0.1 mag
NSVS 19229378	318(6)	LPV	+	0.7	447	3092	
AK Sct	189(4)	LPV	+	0.8	348	3119	
NSVS 19501853	319(2)	M	+	1.4	1476	3191	
ASAS J135447-1640.8	0.37225(4)	EW		0.8	480	3183	
ASAS J195215+1943.2	0.6169(6)	E:	+	0.5	277	2386	
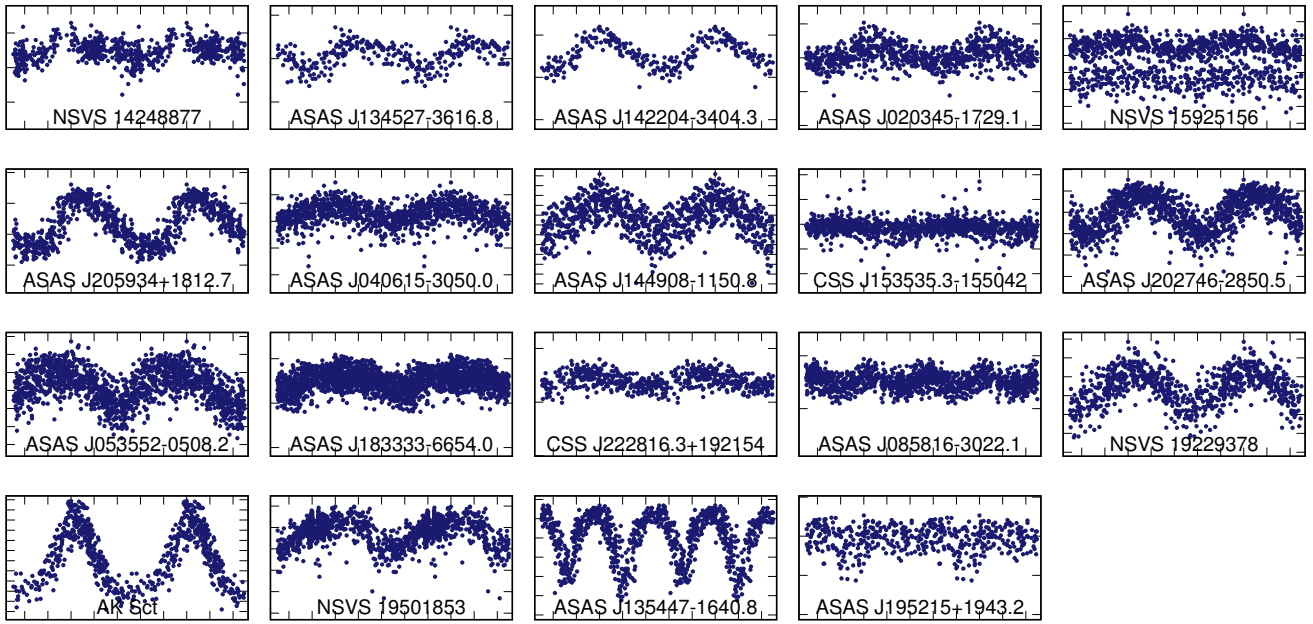
Fig. B.1. Stars misclassified as RRab stars in VSX identified in our sample. Data are phased with periods from Table B.5. The scales of axes are the same as in Fig. 3.