

GALAXY CLUSTERS DISCOVERED VIA THE SUNYAEV–ZEL’DOVICH EFFECT IN THE 2500-SQUARE-DEGREE SPT-SZ SURVEY

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ABSTRACT

We present a catalog of galaxy clusters selected via their Sunyaev–Zel’dovich (SZ) effect signature from 2500 deg² of South Pole Telescope (SPT) data. This work represents the complete sample of clusters detected at high significance in the 2500 deg² SPT-SZ survey, which was completed in 2011. A total of 677 (409) cluster candidates are identified above a signal-to-noise threshold of $\xi = 4.5$ (5.0). Ground- and space-based optical and near-infrared (NIR) imaging confirms overdensities of similarly colored galaxies in the direction of 516 (or 76%) of the $\xi > 4.5$ candidates and 387 (or 95%) of the $\xi > 5$ candidates; the measured purity is consistent with expectations from

simulations. Of these confirmed clusters, 415 were first identified in SPT data, including 251 new discoveries reported in this work. We estimate photometric redshifts for all candidates with identified optical and/or NIR counterparts; we additionally report redshifts derived from spectroscopic observations for 141 of these systems. The mass threshold of the catalog is roughly independent of redshift above $z \sim 0.25$ leading to a sample of massive clusters that extends to high redshift. The median mass of the sample is $M_{500c}(\rho_{\text{crit}}) \sim 3.5 \times 10^{14} M_{\odot} h_{70}^{-1}$, the median redshift is $z_{\text{med}} = 0.55$, and the highest-redshift systems are at $z > 1.4$. The combination of large redshift extent, clean selection, and high typical mass makes this cluster sample of particular interest for cosmological analyses and studies of cluster formation and evolution.

Key words: cosmology: observations – galaxies: clusters: individual – large-scale structure of universe

Supporting material: machine-readable table

1. INTRODUCTION

Galaxy clusters are the largest collapsed objects in the universe, and their abundance is exponentially sensitive to the conditions and processes that govern the cosmological growth of structure (see Allen et al. 2011, and references therein). In particular, measurements of cluster abundance have the power to constrain the matter density, the amplitude and shape of the matter power spectrum, and the sum of the neutrino masses (Wang & Steinhardt 1998; Wang et al. 2005; Lesgourgues & Pastor 2006), and to test models of the cosmic acceleration (Wang & Steinhardt 1998; Haiman et al. 2001; Weinberg et al. 2013) in ways that complement constraints from other observational methods. Clusters are also unique laboratories in which to characterize the interplay between gravitational and astrophysical processes (see Kravtsov & Borgani 2012 for a recent review). Well-defined cluster samples over a broad redshift range are critical for such studies.

Large samples of clusters were first compiled from optical and infrared data sets, in which clusters are identified as overdensities of galaxies. Clusters identified in this manner typically contain of tens to thousands of galaxies (e.g., Abell 1958; Gladders & Yee 2000; Koester et al. 2007; Eisenhardt et al. 2008; Wen et al. 2012; Rykoff et al. 2014). Clusters of galaxies are also bright sources of extended X-ray emission: the majority of the baryonic matter in clusters (70%–95%) is found in the intracluster medium (ICM) in the form of diffuse gas that has been heated by adiabatic compression and shocks to X-ray emitting temperatures of 10^7 – 10^8 K (see review by Voit 2005). The observational expense of detecting high-redshift systems currently limits the size of X-ray samples; however, the tight correlation between ICM observables and the underlying cluster mass enables even modest samples of X-ray-selected systems to place competitive constraints on cosmological models (e.g., Vikhlinin et al. 2009; Mantz et al. 2010).

Over the past decade, clusters have begun to be identified via their signature in the millimeter-wave sky. As photons from the cosmic microwave background (CMB) pass through a galaxy cluster, roughly 1% of the photons will inverse-Compton scatter off the energetic electrons in the ICM. This imparts a characteristic spectral distortion to the CMB, known as the thermal Sunyaev–Zel’dovich (SZ) effect (Sunyaev & Zel’dovich 1972). The observed temperature difference, ΔT , relative to the mean CMB temperature, T_{CMB} , is

$$\begin{aligned} \Delta T &= T_{\text{CMB}} f_{\text{SZ}}(x) \int n_e \frac{k_B T_e}{m_e c^2} \sigma_T dl \\ &\equiv T_{\text{CMB}} f_{\text{SZ}}(x) y_{\text{SZ}}, \end{aligned} \quad (1)$$

where the integral is along the line of sight, $x \equiv h\nu/k_B T_{\text{CMB}}$, k_B is the Boltzmann constant, c the speed of light, n_e the electron density, T_e the electron temperature, σ_T the Thomson cross-section, $f_{\text{SZ}}(x)$ encodes the frequency dependence of the thermal SZ effect:

$$f_{\text{SZ}}(x) = \left(x \frac{e^x + 1}{e^x - 1} - 4 \right) (1 + \delta_{\text{rc}}) \quad (2)$$

(Sunyaev & Zel’dovich 1980), and y_{SZ} is the Compton y -parameter. The term δ_{rc} represents relativistic corrections (Nozawa et al. 2000), which become important at $T_e \gtrsim 8$ keV. This frequency dependence leads to a decrement of observed photons (relative to a blackbody at T_{CMB}) below the thermal SZ null frequency of ~ 220 GHz and a corresponding increment of photons above this frequency. The surface brightness of the thermal SZ effect is independent of redshift, and the integrated thermal SZ signal is expected to be a low-scatter proxy for the cluster mass, as it is proportional to the total thermal energy of the ICM (Motl et al. 2005). These properties make cluster samples produced by SZ surveys attractive for cosmological analyses (Carlstrom et al. 2002).

The observed temperature distortions in the CMB caused by the thermal SZ effect are small, typically on the level of hundreds of μK_{CMB} for the most massive clusters.⁴⁶ The development over the past decade of high-sensitivity bolometric cameras has enabled the wide and deep surveys required to find these rare systems via the SZ effect. The first discovery of a previously unknown cluster through its SZ signature was published in 2009 (Staniszewski et al. 2009); today such discoveries have become routine, with catalogs produced by the South Pole Telescope (SPT), Atacama Cosmology Telescope (ACT), and *Planck* teams containing tens to hundreds of massive clusters out to $z \sim 1.5$ (Reichardt et al. 2013; Hasselfield et al. 2013; Planck Collaboration et al. 2014a).

In this paper we present a cluster catalog extracted from the full 2500 deg² SPT-SZ survey. This sample consists of 677 galaxy cluster candidates detected at SPT statistical significance $\xi > 4.5$, of which 516 (76%) have been confirmed as clusters via dedicated optical and near-infrared (NIR) follow-up imaging. Follow-up imaging was obtained for all 530 candidates detected above $\xi = 4.7$ and 119 (81%) of the remaining 147 candidates down to $\xi = 4.5$. For all confirmed clusters, we report photometric redshifts—or, where available, spectroscopic redshifts—and estimated masses. Of the confirmed clusters, 251 (49%) are reported for the first time in this work. Masses

⁴⁶ Throughout this work, noise units and amplitudes expressed in terms of K_{CMB} refer to the equivalent deviations in temperature from a 2.73 K blackbody required to produce the observed signals.

Table 1
The Fields Observed by the South Pole Telescope between 2008 and 2011

Name	R.A. ($^{\circ}$)	δ ($^{\circ}$)	Area (deg 2)	σ_{95} (μ K-arcmin)	σ_{150} (μ K-arcmin)	σ_{220} (μ K-arcmin)	Survey Year
RA5H30DEC-55	82.7	-55.0	82.9	38.2	12.8	37.0	2008, 2011
RA23H30DEC-55	352.5	-55.0	100.2	36.9	11.7	35.0	2008, 2010
RA21HDEC-60	315.0	-60.0	147.6	35.5	15.0	58.1	2009
RA3H30DEC-60	52.5	-60.0	222.6	35.7	15.7	59.0	2009
RA21HDEC-50	315.0	-50.0	190.0	40.7	17.7	65.7	2009
RA4H10DEC-50	62.5	-50.0	155.5	30.9	14.4	59.5	2010
RA0H50DEC-50	12.5	-50.0	156.2	36.8	16.1	64.2	2010
RA2H30DEC-50	37.5	-50.0	155.7	35.1	15.2	58.5	2010
RA1HDEC-60	15.0	-60.0	145.9	34.6	15.6	60.1	2010
RA5H30DEC-45	82.5	-45.0	102.7	39.1	17.7	72.7	2010
RA6H30DEC-55	97.5	-55.0	83.3	35.6	15.7	65.2	2011
RA3H30DEC-42.5	52.5	-42.5	166.8	34.0	15.4	62.4	2011
RA23HDEC-62.5	345.0	-62.5	70.5	35.9	15.8	62.2	2011
RA21HDEC-42.5	315.0	-42.5	111.2	36.7	16.5	67.0	2011
RA1HDEC-42.5	15.0	-42.5	108.6	35.4	15.3	62.9	2011
RA22H30DEC-55	337.5	-55.0	83.6	37.3	16.3	67.2	2011
RA23HDEC-45	345.0	-45.0	204.5	35.0	15.6	64.4	2011
RA6H30DEC-45	97.5	-45.0	102.8	35.8	15.6	67.0	2011
RA6HDEC-62.5	90.0	-62.5	68.7	34.9	15.9	67.9	2011

Notes. Descriptive information for the 19 fields that comprise the 2500 deg 2 SPT-SZ survey. Here we list the field name, center, effective area following point source masking, noise levels at 95, 150, and 220 GHz, and the year the field was imaged. The noise levels are estimated as in Schaffer et al. (2011), using the Gaussian beam approximation.

are computed using the framework developed for Reichardt et al. (2013, hereafter R13); we report masses using the best-fit ξ -mass relation for a fixed flat Λ CDM cosmology with $\Omega_m = 0.3$, $h = 0.7$ and $\sigma_8 = 0.8$. A detailed cosmological analysis incorporating new information from follow-up X-ray observations will be presented in T. de Haan et al. (in preparation, hereafter referred to as dH15).

This paper is organized as follows. Observations and map-making are described in Section 2. The extraction of galaxy clusters from the maps is detailed in Section 3. Building off previous work presented in Song et al. (2012, hereafter S12), we describe our optical and NIR follow-up campaign in Section 4 and associated confirmation of clusters and measurement of redshifts in Section 5. In Section 6 we present the full sample of galaxy cluster candidates from the 2500 deg 2 SPT-SZ cluster survey and highlight particularly notable clusters, and we conclude in Section 7. All masses are reported in terms of M_{500c} , where M_{500c} is defined as the mass enclosed within a radius at which the average density is 500 times the critical density at the cluster redshift. Selected data reported in this work, as well as future updates to the properties of these clusters, will be hosted at <http://pole.uchicago.edu/public/data/sptsz-clusters>.

2. OBSERVATIONS AND DATA REDUCTION

2.1. Telescope and Observations

The SPT (Carlstrom et al. 2011) is a 10 m diameter telescope located at the National Science Foundation Amundsen-Scott South Pole station in Antarctica. From 2008 to 2011 the telescope was used to conduct the SPT-SZ survey, a survey of ~ 2500 deg 2 of the southern sky at 95, 150, and 220 GHz. The survey covers a contiguous region from 20^h to 7^h in right ascension (R.A.) and -65 to -40° in declination (see, e.g., Figure 1 in Story et al. 2013) and was mapped to depths of approximately 40, 18, and 70 μ K-arcmin at 95, 150, and 220 GHz respectively.

The telescope was designed for high-resolution measurements of the CMB, with particular attention to the science goal of discovering high-redshift galaxy clusters through the SZ effect. The large primary mirror leads to beam sizes of roughly 1'6, 1'1, and 1'0 at 95, 150, and 220 GHz. Beams of this scale are well matched to the expected angular size of massive clusters at high redshift. For a non-relativistic thermal SZ spectrum, the centers of the measured 95 and 150 GHz bands are at 97.6 GHz and 152.9 GHz, while the 220 GHz band is approximately at the thermal SZ null.

The 2500 deg 2 SPT-SZ survey was not observed as one contiguous patch; rather it was broken into 19 subregions, or fields, individually scanned to survey depth. These fields range in size from ~ 70 to 250 deg 2 , and their borders overlap slightly ($\sim 4\%$). The total area searched for clusters over all fields, after correcting for point-source masking (see Section 3) is 2365 deg 2 .

The majority of fields were observed in an identical fashion: the telescope was scanned back and forth across the width of the field in azimuth and then stepped in elevation; the scan and step procedure was repeated until the full field had been imaged. The process, which constitutes a single *observation* of the field, took between 0.5 and 2.5 hr, depending on the field size and the elevation step. Each field was imaged in this fashion at least 200 times, and final maps are the sum of all individual observations. One field, RA21HDEC-50,⁴⁷ was observed with two different strategies: roughly one-third of the data were obtained with the strategy described above, while the remainder of the observations were conducted by scanning the telescope in elevation at a series of fixed azimuth angles. In Table 1 we have listed the name, location, area, and depths at 95, 150, and 220 GHz for each field, as well as the year in which the field was observed. The depth estimates are obtained as in Schaffer et al. (2011), using the Gaussian beam approximation described in that work.

⁴⁷ Fields are named via the R.A. and declination of their centers.

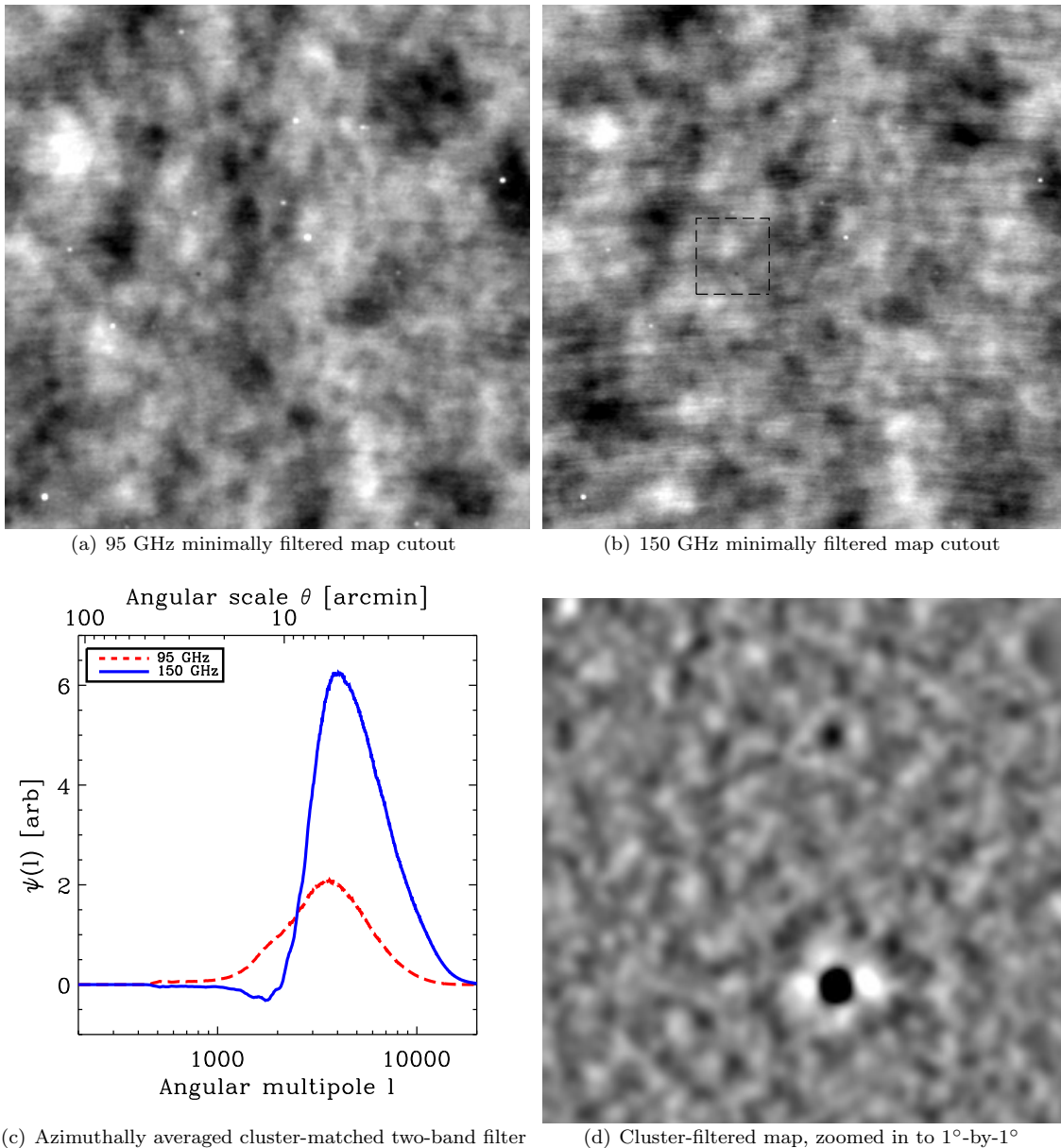


Figure 1. Visual representation of the SPT-SZ data and matched filtering process described in Sections 2 and 3. Panels (a) and (b) show 6° by 6° cutouts of 95 and 150 GHz maps from the RA21HDEC–60 field; the displayed temperature range is $\pm 300 \mu\text{K}$. These maps are made from data that have been only minimally filtered (scan-direction high-pass filter at $l \sim 50$) and show the main features of SPT-SZ survey data: large-scale primary CMB fluctuations, emissive point sources, and SZ decrements from galaxy clusters. Panel (c) shows the azimuthally averaged spatial-spectral filter optimized for detection of $\theta_c = 0.25$ clusters, with the red-dashed (blue-solid) curves showing the Fourier-domain coefficients for the 95 (150) GHz data. Panel (d) shows a zoomed-in view of the 1°-by-1° area delineated by the dashed box in panel (b) after the spatial-spectral filter has been applied. This map is in units of signal-to-noise, and the displayed range is $-5 < S/N < 5$. Visible in this panel are the $\xi = 22.2$, $z = 1.132$ cluster SPT-CL J2106–5844 and the $\xi = 4.6$, optically unconfirmed candidate SPT-CL J2106–5820.

The top two panels of Figure 1 show 6°-by-6° cutouts of 95 and 150 GHz maps from the RA21HDEC–60 field. The maps shown here have been very minimally filtered (high-passed at an angular multipole value of roughly $l = 50$ in the scan direction) and show the main signal components present in SPT-SZ survey data, namely the large-scale primary CMB fluctuations, emissive point sources, and SZ decrements from galaxy clusters. The data described in the next section and used as input to the cluster finding pipeline are more strongly filtered in the time domain before being processed into maps.

2.2. Map Making

A series of operations are performed to convert the raw data from the field observations to calibrated maps. The map-making

process in this work is almost identical to that in R13; the most significant change lies in the treatment of the RA5H30DEC–55 and RA23H30DEC–55 fields. These two fields, originally observed in 2008, were re-observed in 2010 (RA23H30DEC–55) or 2011 (RA5H30DEC–55) in order to add coverage at 95 GHz (the array fielded in 2008 did not have enough high-quality 95 GHz detectors to produce survey-depth data). These additional observations also resulted in deeper final maps at 150 GHz (see Table 1). For these two fields, R13 analyzed the 150 GHz maps described in Vanderlinde et al. (2010). The filtering and calibration for these two fields were treated differently from the other fields. In this analysis we use the full, two-season data (including the 95 GHz data), and the treatment of filtering and calibration is uniform across the survey.

We briefly summarize the map-making process here; these steps are described in detail in Schaffer et al. (2011).

1. For each observation, time-ordered detector data are first notch-filtered to remove sensitivity to the pulse-tube cooler of the cryostat housing the SPT-SZ receiver. Which detectors' data to include in mapmaking is then determined using a series of cuts based upon noise performance and response to both a chopped thermal source and on-sky sources. Following these cuts the array is “flat-fielded” by adjusting each detector’s data according to its response relative to other detectors of the same frequency.
2. In every scan across the field, each detector’s data is high-pass filtered by removing the best-fit Legendre polynomial (with the polynomial order depending on the length of the scan) and a series of sine and cosine modes. The resulting filter has an effective cutoff frequency corresponding to angular multipole $l = 400$ (roughly 1/2 degree scales) in the scan direction. This filtering step removes large-scale noise from the atmosphere and low-frequency noise from the readout. To further reduce atmospheric contamination, the mean signal from detectors in a single wedge⁴⁸ is subtracted from the data of all detectors in that wedge. This common-mode subtraction acts as an isotropic high-pass filter with a cutoff at approximately $l = 500$. Bright point sources detected at $>5\sigma$ (~ 6 mJy at nominal survey depth) in 150 GHz data are masked in both of these filtering steps.
3. Following filtering, the telescope pointing model is used to project the data onto two-dimensional maps. For this analysis we use the Sanson–Flamsteed projection (Calabretta & Greisen 2002) in which pixel rows in the map correspond to constant elevation. As (for the majority of observations) the telescope also scans at constant elevation, this projection simplifies the characterization of the applied filtering at the cost of slight shape distortions at the map edges.
4. Individual maps, weighted by their noise properties at $1500 < l < 4500$, are then coadded to produce the final maps. Maps with anomalously high or low weights or noise are not included in the coadd.
5. A calibration factor based on observations of the galactic H II region RCW38 is applied to provide the absolute temperature calibration for the maps (Staniszewski et al. 2009). We have repeated the cluster-finding procedure described in Section 3 using maps with a CMB-based calibration and find negligible differences in the detection significances of galaxy clusters.

3. CLUSTER EXTRACTION AND MM-WAVE CHARACTERIZATION

In this section, we provide a summary of the process by which galaxy cluster candidates are identified and characterized in the SPT survey data. This procedure is almost identical to that used in recent SPT publications; readers are referred to Williamson et al. (2011) and R13 in particular for more details. The small differences between R13 and this analysis are discussed in Section 6.3.

3.1. Cluster Extraction

As described in Section 2, the SPT-SZ survey fields are observed at 3 frequency bands centered at approximately 95, 150, and 220 GHz. These maps contain signal from a range of astrophysical sources. For the purposes of this analysis, we characterize the observed temperature, T , in the maps at frequency ν_i and location \mathbf{x} by

$$T(\mathbf{x}, \nu_i) = B(\mathbf{x}, \nu_i) * [f_{\text{SZ}}(\nu_i)T_{\text{CMB}}y_{\text{SZ}}(\mathbf{x}) + n_{\text{astro}}(\mathbf{x}, \nu_i)] + n_{\text{noise}}(\mathbf{x}, \nu_i). \quad (3)$$

Here B encompasses the effects of the beam and applied filtering; the expected thermal SZ signal is given by the product of the frequency dependent term f_{SZ} , the CMB temperature T_{CMB} , and the Compton- y parameter y_{SZ} ; n_{astro} encompasses astrophysical signals—all of which are modeled here as Gaussian noise—and n_{noise} corresponds to instrumental and residual atmospheric noise not removed by the filtering discussed in Section 2. For SPT maps, n_{astro} primarily consists of lensed primary CMB fluctuations, kinetic and thermal SZ from the clusters below the SPT detection threshold, and dusty extragalactic sources; radio sources below the SPT detection threshold contribute negligibly to the maps. As in previous work, we model these noise terms based upon recent SPT power spectrum constraints (Keisler et al. 2011; Shirokoff et al. 2011).

Given the known spatial and spectral characteristics of galaxy clusters as well as the sources of noise in the maps, we construct a filter designed to maximize our sensitivity to galaxy clusters (Melin et al. 2006). This Fourier-domain filter takes the form:

$$\psi(\mathbf{l}, \nu_i) = \sigma_{\psi}^{-2} \sum_j \mathbf{N}_{ij}^{-1}(\mathbf{l}) f_{\text{SZ}}(\nu_j) S_{\text{filt}}(\mathbf{l}, \nu_j), \quad (4)$$

where the predicted variance in the filtered map, σ_{ψ}^{-2} , is given by

$$\sigma_{\psi}^{-2} = \int d^2l \sum_{i,j} f_{\text{SZ}}(\nu_i) S_{\text{filt}}(\mathbf{l}, \nu_i) \mathbf{N}_{ij}^{-1}(\mathbf{l}) \times f_{\text{SZ}}(\nu_j) S_{\text{filt}}(\mathbf{l}, \nu_j), \quad (5)$$

\mathbf{N} is the Fourier-domain version of the band–band, pixel–pixel covariance matrix, and S_{filt} is the Fourier transform of the real space cluster template convolved with $B(\mathbf{x}, \nu_i)$. We use a projected isothermal β -model (Cavaliere & Fusco-Femiano 1976) with β fixed to 1 as our source template:

$$\Delta T = \Delta T_0 (1 + \theta^2/\theta_c^2)^{-1}, \quad (6)$$

where ΔT_0 is the normalization, θ is the angular separation from the cluster center, and θ_c is the core radius. As discussed in Vanderlinde et al. (2010), given the spatial resolution of the SPT this simple profile is adequate for our purposes: no improvement in the detection of clusters is seen using more sophisticated models (e.g., Nagai et al. 2007; Arnaud et al. 2010).

A series of profiles with evenly spaced core radii ranging from 0.25 to 3' is used to construct 12 matched filters. Azimuthal averages of the 0.25 matched-filter coefficients for 95 and 150 GHz are shown as a function of ℓ in the bottom left panel of Figure 1. After the application of these filters to the 95–150 GHz maps (the relative noise levels in the 220 GHz maps are too high to significantly improve cluster detection, so these data are omitted here), cluster candidates are extracted via

⁴⁸ The modular SPT-SZ receiver consists of 6 independent sub-arrays or “wedges” of 160 bolometers that operate at a single frequency.

a peak detection algorithm similar to the SExtractor routine (Bertin & Arnouts 1996). We record the location and maximum detection significance across all filter scales, ξ , for all cluster candidates with $\xi \geq 4.5$. The bottom-right panel of Figure 1 shows the result of applying the 0/25 matched filter to 95 and 150 GHz maps of the RA21HDEC–60 field.

We take two steps to reduce the number of spurious sources created by filter artifacts (in particular decrements created by the filter “ringing” around strong sources). First, prior to filtering the maps we mask a $4'$ region around all point sources detected above 5σ in 150 GHz maps optimized for point source detection. Following filtering we additionally exclude cluster candidates detected within $8'$ of these emissive sources. This step removes 163 deg^2 from the survey. We perform a separate cluster-finding analysis on these masked regions and report on those detections (which we do not include in the official catalog) in Section 6.2. We also find it necessary to veto candidates around the strongest cluster detections: we exclude candidates within a $10'$ region around $\xi > 20$ detections. This final cut removes less than 1 deg^2 from the entire survey.

3.2. Integrated Comptonization

For every cluster candidate we measure the integrated Comptonization, Y_{SZ} , within a 0.75 radius aperture. For a projected isothermal β -model, Y_{SZ} is defined as

$$Y_{\text{SZ}}^{0.75} = 2\pi \int_0^{0.75} y_0 (1 + \theta^2/\theta_c^2)^{-1} \theta d\theta, \quad (7)$$

where y_0 is the peak Comptonization. The aperture is slightly smaller than the $1'$ radius utilized in R13: as demonstrated in Saliwanchik et al. (2015), at the resolution and noise levels of the SPT-SZ survey, the scatter of Y_{SZ} at fixed cluster mass is minimized at $\theta = 0.75$. This scatter, measured to be $27 \pm 2\%$, is comparable to the scatter in the relation between SPT detection significance and mass that is used to estimate cluster masses in R13 and other SPT cluster publications.

Y_{SZ} is computed using the same procedure as in R13. Briefly, for every cluster candidate, the likelihood of the observed two-band SZ signal given the model of Equation (6) is estimated using a simple gridded parameter search. For every point in the four-parameter space of y_0 , θ_c , x position, and y position, the value of $Y_{\text{SZ}}^{0.75}$ is calculated using Equation (7), and the best-fit value and 1σ constraints are estimated from the (one-dimensional) $Y_{\text{SZ}}^{0.75}$ posterior distribution. Effective step-function priors are placed on the values of θ_c , x , and y by restricting the parameter grid for each cluster such that x and y are within 1.5 of the best matched-filter position for that cluster (we note that the absolute difference between the best-fit and matched-filter positions is <0.75 for 97% of the cluster candidates) and that the physical core radius of each cluster is $50 \text{ kpc} \leq r_{\text{core}} \leq 1 \text{ Mpc}$. For unconfirmed candidates we use a redshift of $z = 1.5$ to set the physical scale. The priors on r_{core} are motivated by the known mass distribution of SPT-selected clusters and the mass-concentration relation measured by, e.g., Mandelbaum et al. (2008).

3.3. Contamination

Simulations are used to estimate the level of contamination in the cluster catalog by false detections from instrumental noise and non-cluster astrophysical signals. The simulations are similar to those used in R13; we provide a brief summary here.

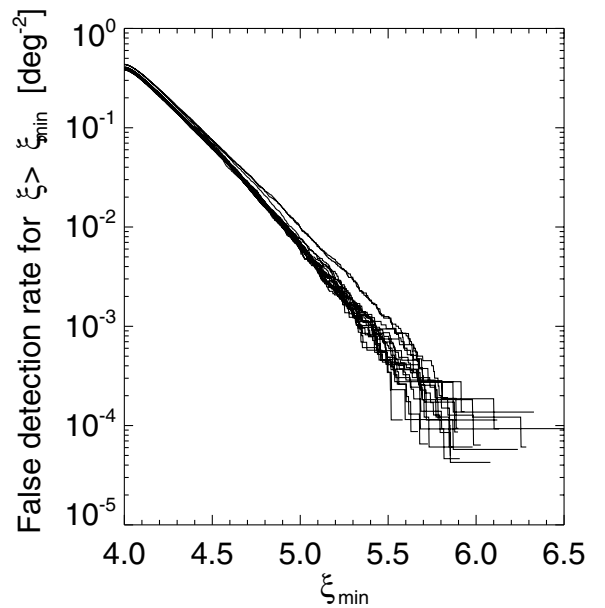


Figure 2. Simulated false-detection rate for each of the 19 fields in the SPT-SZ survey. Plotted is the cumulative density of false detections above a detection significance, ξ_{min} . We expect 18.5 false detections at $\xi \geq 5$ and 172 at $\xi \geq 4.5$ for the entire survey.

For each of the 19 SPT fields we create 100 simulated sky maps composed of contributions from the CMB, emissive sources, and noise (note that we do not include a thermal SZ contribution when quantifying the expected number of false detections). The CMB component is modeled as a Gaussian random field based on the best-fit WMAP7 + SPT lensed Λ CDM model (Komatsu et al. 2011; Keisler et al. 2011), and the point source model contains contributions from radio sources and dusty star-forming galaxies (DSFGs): all source contributions are modeled as Gaussian fields. The simulated radio population follows the results of De Zotti et al. (2005), Vieira et al. (2010), and Reichardt et al. (2012). We assume 100% correlation between the bands, a spectral index of $\alpha = -0.53$ and, at 150 GHz, an amplitude of $D_l = l(l+1)C_l/2\pi = 1.28 \mu\text{K}^2$ at $l = 3000$. The amplitudes and spectral indices of the DSFG contributions are also constrained by recent SPT measurements (Reichardt et al. 2012). At 150 GHz and $l = 3000$ the Poisson contribution has amplitude $D_l = 7.54 \mu\text{K}^2$ and the clustered contribution $D_l = 6.25 \mu\text{K}^2$; we use $\alpha = 3.6$ for both contributions. For each simulated map, we create noise realizations using jackknife noise maps (e.g., Sayers et al. 2009). The noise maps are created by randomly multiplying half of the observations of a field by -1 and then coadding the entire set of observations. This is a change from the R13 simulations which assumed stationary Gaussian noise.

To estimate the expected number of false detections, the cluster-detection algorithm is run (with point source masks and apodization matching the real data) on these cluster-free simulated maps. In Figure 2 we plot the expected rate of false detections in each individual SPT field. Two fields have slightly higher false-detection rates owing to the inclusion of boundary regions with uneven coverage in the area searched for clusters. In total, across all fields, we expect 172 false detections at $\xi > 4.5$ and 18.5 at $\xi > 5$. We return to the question of false detections again in Section 6 where we compare our expectations to the measured purity of the cluster sample.

Table 2
Optical and Infrared Imagers

Ref. ^a	Site	Telescope	Aperture (m)	Camera	Filters ^b	Field
1	Cerro Tololo	Blanco	4	MOSAIC-II	<i>griz</i>	36' × 36'
2	Las Campanas	Magellan/Baade	6.5	IMACS f/2	<i>griz</i>	27.4 × 27.4
3 ^c	Las Campanas	Magellan/Clay	6.5	LDSS3	<i>griz</i>	8.3 diam. circle
4 ^d	Las Campanas	Magellan/Clay	6.5	Megacam	<i>gri</i>	25' × 25'
5	Las Campanas	Swope	1	SITe3	<i>BVRI</i>	14.8 × 22.8
6	La Silla	MPG/ESO	2.2	WFI	<i>BVRI</i>	34' × 33'
7	La Silla	New Technology Telescope	3.6	EFOSC2	<i>griz</i>	4.1 × 4.1
8	Cerro Tololo	Blanco	4	NEWFIRM	<i>K_s</i>	28' × 28'
9	Las Campanas	Magellan/Baade	6.5	FourStar	<i>J, H, K_s</i>	10.8 × 10.8
10	Satellite	<i>Spitzer Space Telescope</i>	0.85	IRAC	3.6 μm, 4.5 μm	5.2 × 5.2
...	Satellite	<i>Wide-field Infrared Survey Explorer</i>	0.40	...	<i>W1, W2</i>	47' × 47'

Notes. Optical and infrared cameras used in SPT follow-up observations.

^a Shorthand alias used in Table 4.

^b Not all filters were used to image every cluster.

^c <http://www.lco.cl/telescopes-information/magellan/instruments/ldss-3>

^d Megacam data were acquired for a large follow-up weak-lensing program.

4. FOLLOW-UP OBSERVATIONS

We use optical and in some cases NIR imaging to confirm candidates as clusters and to obtain redshifts for confirmed systems. In this section, we briefly describe our follow-up strategy and the reduction of imaging data. Follow-up imaging was obtained for all 530 candidates detected above $\xi = 4.7$ and 119 of the remaining 147 candidates down to $\xi = 4.5$; the remaining 28 low-significance candidates were not imaged owing to limited follow-up resources. The procedures discussed here closely follow those presented in S12.

4.1. Follow-up Strategy

Our candidate follow-up strategy evolved over the course of the SPT-SZ survey. Initially, clusters were confirmed using preexisting imaging from the Blanco Cosmology Survey (BCS; Menanteau et al. 2010b; Desai et al. 2012; Bleem et al. 2015, see Section 4.3) as well as targeted imaging on the Blanco/MOSAIC-II⁴⁹, and Magellan/IMACS imagers (Dressler et al. 2006) (see High et al. 2010 for more details). As the candidate sample size grew, we adjusted our imaging strategy to effectively incorporate a range of small- and large-aperture telescopes. Our resulting strategy is as follows.

1. All cluster candidates are “pre-screened” using imaging data from the Digitized Sky Survey (DSS)⁵⁰ to determine if a cluster candidate lies at low redshift. We find the DSS images are generally sufficient to allow identification of the optical counterparts of SPT clusters out to redshift $z = 0.5$, with a tail extending to $z = 0.7$. Roughly 50% of confirmed SPT clusters are identifiable in DSS data; candidates not apparent in the DSS are classified as high-redshift targets.
2. We observe potential low-redshift candidates at 1–2 m class facilities. If the candidate is not confirmed in these data, it is reclassified as a potential high-redshift system.
3. High-redshift targets are imaged on larger aperture (4–6.5 m) telescopes. The majority of these cluster candidates with $\xi \geq 4.8$ and many candidates with $4.5 \leq \xi < 4.8$ have also been imaged in the NIR from ground-

space-based facilities. This threshold was $\xi = 4.5$ in S12; the increased threshold here is due to limited NIR resources.

4. Following the release of the *Wide-field Infrared Survey Explorer* (WISE) all-sky catalogs (Wright et al. 2010), observations of $\xi < 4.8$ candidates were weighted toward those candidates with lower values of the NIR “blank-field” statistic presented in S12. This statistic quantifies the significance of any overdensity of galaxies at the candidate’s location compared to random fields; lower values correspond to more dense regions. As such, the follow-up of these lower significance candidates is biased to maximize the number of confirmed clusters.

In Table 2 we list the facilities and instruments used in our follow-up campaign. We assign each telescope/instrument combination a numerical alias which we use in Table 4 to identify the source of redshift information.

4.2. Targeted Observations and Data Reduction

4.2.1. Optical Data

Our strategy for targeted optical follow-up varied with the aperture of the follow-up telescope. In this section, we first describe our strategy for observing candidates we expect to be at low redshift with 1 m–2 m class telescopes; we then move on to our strategy for observing likely high-redshift candidates with larger telescopes. Note that in this section we quote depths relative to the apparent magnitude of L^* galaxies: our model for the redshift evolution of these galaxies is described in Section 5.

In order to rapidly image a large number of systems, we adopt a minimalist approach for low-redshift candidates observed on the 1 m Swope telescope. Based on an initial “by-eye” redshift estimate from the DSS screening step,⁵¹ we choose a pair of filters (*BV, VR, RI*) expected to straddle the 4000 Å break. Three filters are used when required to avoid redshift degeneracies. Candidates are imaged to depths sufficient for robust estimation of red-sequence redshifts: we require detection of $0.4L^*$ red-sequence galaxies at 8σ in the redder filter and 5σ in the bluer filter. A second round of deeper imaging is obtained for

⁴⁹ <http://www.ctio.noao.edu/mosaic/>

⁵⁰ The DSS is a digitization of the photographic sky survey conducted with the Palomar and UK Schmidt telescopes, <http://archive.stsci.edu/dss/>.

⁵¹ These crude redshifts are based on a combination of the color, brightness, and angular size of identified cluster galaxies. The uncertainty on these estimates is $\sigma_z \sim 0.1$, with large outliers owing to variable DSS image quality.

systems with significantly underestimated DSS redshifts. Non-confirmed candidates and clusters at higher redshift ($z > 0.7$) are re-observed on larger class telescopes and/or with NIR imagers.

Low-redshift candidates are also observed on the MPG/ESO 2.2 m telescope using the Wield Field Imager (WFI; Baade et al. 1999). These observations are deeper than those acquired on Swope as they were designed to also enable studies of the galaxy populations of these clusters (e.g., Zenteno et al. 2011). Based on its preliminary DSS redshift, each candidate is imaged in three filters (*BVR* or *VRI*), to depths sufficient to detect $0.4L^*$ galaxies at 10σ in the bands straddling the 4000 Å break. The third band is used for photometric calibration (see below). Second-pass imaging is obtained as necessary to adjust for imperfect initial redshift estimates. As with our Swope program, non-confirmed candidates and clusters at higher redshift ($z > 0.75$) are re-observed on larger class telescopes and/or with NIR imagers.

As described in S12, we also adopt a two-pass strategy for observations on 4–6.5 m class telescopes. Candidates are first imaged in the *g*, *r*, *z* bands (or *g*, *r* and *i* bands early in the follow-up campaign) to depths sufficient to detect $0.4L^*$ galaxies at $z = 0.75$ at 5σ in the redder bands. The *g*-band data is used for photometric calibration. Following this first-pass imaging, all non-confirmed candidates at $\xi \geq 4.8$ (and a subsample of non-confirmed candidates below this threshold) were further imaged in the *z* and *i* (or *r*) bands to extend this redshift range to $z = 0.9$ and/or were imaged in the NIR, as described below.

With the exception of images from Magellan/Megacam (McLeod et al. 2006), all optical images were reduced with the PHOTPIPE pipeline (Rest et al. 2005; Garg et al. 2007; Miknaitis et al. 2007). Megacam images were reduced using the Smithsonian Astrophysical Observatory Megacam reduction pipeline. The PHOTPIPE reduction process as applied to SPT clusters is explained in High et al. (2010) and the Megacam pipeline in High et al. (2012). All reductions include the standard CCD image processing steps of masking bad or saturated pixels, applying crosstalk and overscan corrections, debiasing, flat-fielding, correcting for scattered light via illumination corrections, and—where necessary in the redder bands—defringing. Cosmic rays are also removed from the Megacam images. Images are coadded using the SWarp algorithm (Bertin et al. 2002), and astrometry is tied to the Two Micron All Sky Survey (2MASS) catalog (Skrutskie et al. 2006).

Sources are detected using the SExtractor algorithm (v 2.8.6) in dual-image mode; we use the deepest images with respect to red-sequence galaxies as the detection images. As in previous works, photometry is calibrated using stellar locus regression (SLR; High et al. 2009) with absolute calibration derived using stars in the 2MASS catalog.

There are a few modifications to the calibration process discussed in previous SPT publications. For clusters only imaged in two bands at Swope, we combine *J*-band data from 2MASS with the optical data to create the two colors required for calibration via SLR. The SLR calibration for candidates imaged on the NTT was somewhat challenging owing to the small field of view ($4'.1 \times 4'.1$) of the EFOSC2 imager (Buzzoni et al. 1984). These calibrations were performed with fewer stars than the other imaging; for a few fields we jointly calibrated the data with other observations from the same night. We increase the expected uncertainty on the color calibration for these systems to 5% (compared to the typical 2%–3% observed with SLR, see High et al. 2009; Bleem et al. 2015) and include this extra scatter in our estimates of redshift uncertainties. Finally, for

a small subset of clusters located in the wings of the Large Magellanic Cloud (LMC), we restrict our fitting to stars with counterparts in the 2MASS catalog to enable convergence of the regression algorithm.⁵²

4.2.2. Near-infrared Data

Spitzer/IRAC imaging (Fazio et al. 2004) at $3.6 \mu\text{m}$ and $4.5 \mu\text{m}$ is obtained for the majority of high-redshift SPT cluster candidates at $\xi \geq 4.8$ and a subsample of systems at $4.5 \leq \xi \leq 4.8$. In total 276 candidates (241 candidates at $\xi > 4.8$) were observed as part of our *Spitzer* follow-up program.⁵³ These *Spitzer* data play a crucial role in confirming and determining the redshift of clusters at $z > 0.8$. Candidates are imaged in 8×100 s and 6×30 s dithered exposures at 3.6 and $4.5 \mu\text{m}$, respectively; the resulting coadded images are sufficient for the detection in the $3.6 \mu\text{m}$ band of $z = 1.5$ $0.4L^*$ galaxies at 10σ . These observations are reduced following the methodology of Ashby et al. (2009). Briefly, these reductions correct for column pull-down, mosaic the individual exposures, resample the images to $0''.86$ pixels (half the solid angle of the native IRAC pixels), and reject cosmic rays.

Ground-based NIR imaging of some candidates was acquired with the NEWFIRM imager (Autry et al. 2003) at the CTIO 4 m Blanco telescope and the FourStar imager (Persson et al. 2013) on the Magellan Baade 6.5 m telescope. NEWFIRM data with a target 10σ point source depth of 19 Vega magnitudes in the K_s filter were obtained for 31 candidates during two runs in 2010 November and 2011 July under photometric conditions. Typical observations consisted of 16 point dither patterns, with 6×10 s exposures obtained at each dither position. The data were reduced using the FATBOY pipeline, originally developed for the FLAMINGOS-2 instrument, and modified to work with NEWFIRM data in support of the Infrared Bootes Imaging Survey (Gonzalez et al. 2010). Images were coadded using SCAMP and SWarp (Bertin et al. 2002) and photometry was calibrated to 2MASS.

Additional JHK_s -band imaging was collected with FourStar for 34 candidates during several runs in 2012, 2013, and 2014 in average to good conditions. Several exposures were taken at 9–15 different pointed positions with the coordinates of the cluster centered on either the mosaic or one of the four detectors. The images were flat-fielded using standard IRAF routines; WCS registering and stacking were done using the PHOTPIPE pipeline and were calibrated photometrically to 2MASS.

4.2.3. Spectroscopic Observations

We have also used a variety of facilities to obtain spectroscopic observations of SPT clusters. These observations fall into two categories: small, few-night programs focused primarily on the highest-redshift subset of the SPT-selected clusters (e.g., Brodwin et al. 2010; Foley et al. 2011; Stalder et al. 2013; Bayliss et al. 2014) and longer, multi-semester campaigns. The longer programs include observations of high-redshift systems using FORS2 on the Very Large Telescope (VLT; Appenzeller et al. 1998) and a large survey program on the Gemini-South telescope (NOAO PID 2011A-0034) using GMOS-S

⁵² The stellar envelope of the LMC extends well beyond regions of significant thermal dust emission; we see no evidence of contamination in the SPT millimeter-wave maps.

⁵³ Archival observations with varying wavelength coverage and exposure times are available for an additional 16 (typically low-redshift) SPT systems.

(Hook et al. 2004) that targeted 85 SPT clusters in the redshift interval $0.3 < z < 0.8$. In Ruel et al. (2014) we describe in detail our spectroscopic followup campaign and report spectroscopic redshift measurements for 61 SPT clusters and velocity dispersions for 48 of these clusters. Here we report an additional 34 cluster redshifts using newly obtained data from the Gemini Survey Program. As described below in Section 6, we also search the literature for spectroscopic counterparts of SPT clusters; in total 141 of the clusters in this work have spectroscopically measured redshifts.

4.3. Other Data Sets

In addition to dedicated optical/NIR observations of SPT cluster candidates, we use imaging from three surveys that overlap the SPT footprint: the BCS, the *Spitzer*-South Pole Telescope Deep Field (SSDF; Ashby et al. 2013), and the *WISE* all-sky survey. The BCS is a ~ 80 deg² four-band (g, r, i, z) survey (NOAO large survey program 2005B-0043) with imaging sufficient for cluster confirmation to $z \sim 1$. It is composed of two fields roughly centered at (R.A., decl.) = ($23^{\text{h}}, -55^{\text{d}}$) and ($5^{\text{h}}30^{\text{m}}, -53^{\text{d}}$). These fields roughly overlap with the RA5H30DEC-55 and RA23H30DEC-55 fields, the first fields surveyed by the SPT (see Table 1). We use the reductions presented in Bleem et al. (2015) in this work.

The SSDF, a 94 deg² survey at 3.6 and 4.5 μm , is centered at (R.A., decl.) = ($23^{\text{h}}30, -55^{\text{d}}$). It encompasses a large fraction of the RA23H30DEC-55 field. This survey, one of the largest extragalactic surveys ever conducted with *Spitzer*/IRAC, has imaging sufficient for cluster confirmation and redshift estimation to $z \gtrsim 1.5$.

The *WISE* all-sky survey provides catalogs and images of the entire sky in the W1-W4 bands (3.4–22 μm)⁵⁴; the shorter-wavelength NIR data from *WISE* are sensitive to cluster galaxies out to $z \sim 1.3$ (Gettings et al. 2012; Stanford et al. 2014). As discussed in Section 4.1, we use *WISE* data to prioritize the follow-up of lower significance candidates.

5. CLUSTER CONFIRMATION AND REDSHIFT ESTIMATION

As in previous SPT publications, we deem a candidate to be “confirmed” if we identify an excess of clustered red-sequence galaxies at the SPT location. In this section, we describe our model for the optical and NIR properties of red-sequence galaxies, the process by which we identify excesses of such galaxies at candidate locations, and the estimation of redshifts for confirmed clusters using optical and/or NIR data. Finally, for unconfirmed candidates, we describe our procedure for determining the redshift to which our imaging is sufficient to confirm the candidate as a cluster.

5.1. Red-sequence Model

We create our model for the color–magnitude relation of red-sequence galaxies using the GALAXEV package (Bruzual & Charlot 2003). We model the galaxies as passively evolving, instantaneous-burst stellar populations with a formation redshift of $z = 3$; the stellar populations are generated using the Salpeter initial mass function (Salpeter 1955) and follow the Padova 1994 evolutionary tracks (Fagotto et al. 1994). Metallicities are chosen based upon analytical fits to RCS2 cluster data (B. P. Koester 2013, private communication), and cubic splines are used to

interpolate the discrete output of the code to arbitrary redshifts. We compare our stellar-synthesis $m^*(z)$ model to the Rykoff et al. (2012) model for the maxBCG cluster sample over the redshift range for which that model is valid ($0.05 < z < 0.35$). We find a small (~ 0.2 mag) offset, and we correct our model for this offset. Our red-sequence model is further calibrated to real data using the SPT spectroscopic subsample as described below.

5.2. Identifying Red-sequence Overdensities and Estimating Optical Redshifts

For each cluster candidate, we search for a redshift at which there is a clear excess of galaxies near the candidate position that are consistent with the expected red sequence at that redshift. We typically search for overdensities out to $z \sim 0.7$ ($z \sim 1$) for systems with first-pass (second-pass) optical imaging. At a series of discrete redshifts in this range, we compute the background-subtracted, weighted red-sequence galaxy count in a $2'$ ($3'$ at $z < 0.3$) region around the SPT position. The contribution of each galaxy to the weighted sum is based upon the consistency of the galaxy’s color and magnitude with the red-sequence model at the redshift in question.

To confirm a cluster, we require a significant peak in background-subtracted weighted counts. The preliminary redshift is identified as the location of this peak. In a few instances where this peak is marginally significant (for example, when the cluster is well-detected in one imaging band but poorly in the second owing to incomplete follow-up), we confirm clusters based solely on visual identification of member galaxies and manually select cluster galaxies for redshift estimation.

To further refine the preliminary redshift and to estimate a statistical uncertainty, we next bootstrap resample the galaxies that contribute to the peak. For each bootstrap sample, the estimated redshift is the redshift at which the χ^2 statistic

$$\chi^2 = \sum_{\text{galaxies}} \frac{[\text{Model}(\text{magnitude}, \text{color}, z) - \mathbf{g}]^2}{\text{color error}^2 + \sigma_{\text{rs}}^2} \quad (8)$$

is minimized. Here \mathbf{g} encodes the color and magnitude of the galaxies, and $\sigma_{\text{rs}} = 0.05$ (Koester et al. 2007; Mei et al. 2009) is the intrinsic spread of the red sequence. The reported redshift is the median redshift of 100 bootstrap resamples. The final redshift uncertainties are reported as the statistical error estimated from this bootstrap resampling process (typically small as most estimates are derived from tens of galaxies) added in quadrature with a redshift-dependent scatter determined during the spectroscopic tuning of the red-sequence model, which we now describe.

We first estimate “raw” redshifts using the uncalibrated red-sequence model for 103 clusters with good follow-up data and spectroscopic redshifts in the SPT sample. As follow-up observations span different instruments with different combinations of filters (e.g., subsets of *griz* on IMACS, LDSS3, Megacam, MOSAIC-II, and EFOC2 and *BVRI* on the SITE3 and WFI), we separately calibrate models for each color–magnitude combination used in this analysis. We also calibrate models for the Swope/SITE3 and MPG/ESO WFI data separately as we have not transformed the natural Swope photometry to standard *BVRI* passbands. The large number of clusters with spectroscopic redshifts⁵⁵ enables these independent calibrations.

⁵⁵ Spectroscopic redshifts were measured or identified from the literature for 24 clusters observed with the MPG/ESO telescope, 56 clusters imaged with the Swope telescope, and 100 clusters observed with the larger aperture telescopes. Some clusters with spectroscopic redshifts were observed with multiple instruments to facilitate calibration of red-sequence models.

⁵⁴ <http://wise2.ipac.caltech.edu/docs/release/allsky/>

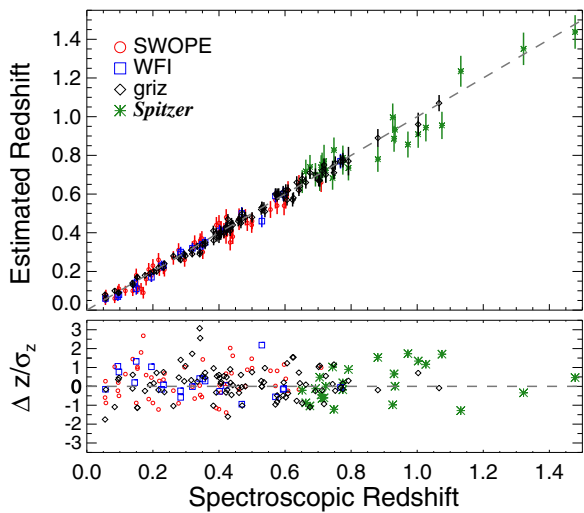


Figure 3. Results of red-sequence model calibration. Top: photometric redshift, z_{est} , vs. spectroscopic redshift, z_{spec} , for 129 spectroscopically confirmed SPT clusters. We plot the aggregate model tunings for Swope/SITe3, MPG/ESO WFI, the larger class telescopes (for which the model was calibrated across instruments as all clusters were observed in the *griz* system), and *Spitzer*/IRAC. Some clusters are plotted multiple times (at most once per model) as they were observed with multiple telescopes to calibrate the various redshift models. Bottom: distribution of redshift residuals $\Delta z/\sigma_z = (z_{\text{spec}} - z_{\text{est}})/\sigma_{z_{\text{est}}}$. The typical redshift uncertainty, σ_z , scales as $\sim 0.013\text{--}0.018(1+z)$ for redshifts estimated using combinations of *griz* filters, $\sim 0.021(1+z)$ for clusters imaged with the MPG/ESO WFI, $\sim 0.025(1+z)$ for Swope/SITe3, and as $\sim 0.035(1+z)$ for redshifts determined using *Spitzer*/IRAC.

We find a linear remapping of model redshifts, z_{model} , to spectroscopic redshifts, z_{spec} ,

$$z_{\text{spec}} = A z_{\text{model}} + B \quad (9)$$

is sufficient for tuning the $g-r$ versus i (or z) relation over the redshift range $z < 0.35$ as well as for tuning the Johnson color–magnitude combinations ($B-V$ versus R , $V-R$ versus R , $R-I$ versus R).

However, as noted in Bleem et al. (2015), we observe large residuals when applying such a first order correction over the broad redshift range sampled by the r,i,z filters. As we expect the remapping from raw to calibrated redshifts to be smoothly varying and monotonic, we use non-linear least squares minimization to fit the z_{model} and z_{spec} relation to a monotonic function. This function is generated using the methodology of Ramsay (1998) where we have chosen sines and cosines as the basis functions and include these functions to the 4th order. As in S12, we estimate the uncertainty in our model calibration by determining the quantity σ_z such that the reduced chi-squared statistic, χ_{red}^2 :

$$\chi_{\text{red}}^2 = \frac{1}{\nu} \sum \frac{(z_{\text{est}} - z_{\text{spec}})^2}{(\delta_z(1+z))^2} = 1, \quad (10)$$

where z_{est} is our calibrated redshift and ν is the number of degrees of freedom. Here the total degrees of freedom are reduced by 2 by the linear rescaling and by 10 for higher order rescaling. We find $\delta_z \sim 0.025$ for Swope/SITe3, $\delta_z \sim 0.021$ for WFI and $\delta_z \sim 0.013\text{--}0.018$ for the *griz*-based redshift models. We plot the results of the redshift tuning in Figure 3.

5.3. Near-infrared Redshifts

We analyze the NIR data as in “Method 1” from S12 (see Section 3.1 in that work) by supplementing the optical with

the JHK_s and IRAC imaging. The IRAC imaging is only used in the cases where the optical imaging is not deep enough to confirm a cluster and measure the redshift. The JHK_s data are used as additional filters to measure at least one color across the 4000 Å break. It should be noted that the *Spitzer*-only [3.6]–[4.5] colors are probing near the peak in the stellar emission, rather than the Balmer break; as such, they are less sensitive to the effects of recent star formation or active galactic nucleus activity which may be more prevalent at high redshift (Brodwin et al. 2013). We have demonstrated with spectroscopic follow-up that this measurement is reliable in the redshift range relevant for confirmation of high- z clusters (Stalder et al. 2013; Bayliss et al. 2014, and see Figure 3). There are three clusters with [3.6]–[4.5] galaxy colors consistent with redshifts greater than 1.5; as our models have not been tested with spectroscopic data in this redshift range, we report a lower limit of $z = 1.5$ for the redshifts of these systems (see Figure 4).

5.4. Redshift Limits

As all of our optical/NIR observations have finite depth, it is not possible to definitively rule out the existence of undiscovered, high-redshift counterparts for our unconfirmed cluster candidates. We instead report for each unconfirmed candidate the highest redshift for which we would have detected the overdensity of red galaxies we require to confirm a cluster. The depth of our follow-up imaging varies among candidates, so this limit is computed individually for each candidate. For every candidate, we determine the redshift for which a $0.4L_*$ red-sequence galaxy matches the measured 10σ magnitude limit of the imaging data. As we require two filters to measure a redshift, we obtain the “redshift limit” from the second deepest of the imaging bands used in the red-sequence overdensity search. A detailed description of this procedure is provided in S12.

6. CLUSTER CATALOG

In Table 4 we present the complete sample of galaxy cluster candidates detected at $\xi \geq 4.5$ in the 2500 deg² SPT-SZ survey. For each candidate we provide the position, integrated Comptonization within an 0.75 -radius aperture (Section 3.2), the candidate detection significance ξ at the filter scale that maximizes detection significance, the value of the β -model core radius θ_c at which ξ is reported, the estimated mass and redshift (spectroscopic where available) for confirmed clusters and redshift limit for unconfirmed candidates. We discuss the estimation of these cluster masses in Section 6.1. In Figure 5 we plot the SZ detection significance versus redshift for each confirmed cluster as well as the redshift distribution of the confirmed cluster sample.

Simulations predict that this catalog should contain only a small number of false detections (see Section 3.3 for details of the simulations), and this prediction is borne out by our optical/NIR follow-up observations (Section 4). Our simulations predict 18.5 false detections above $\xi = 5$ for the full survey—corresponding to a predicted purity of 95% for the 402 $\xi \geq 5$ candidates in our catalog—and 172 false detections in the full $\xi \geq 4.5$ sample—corresponding to a predicted purity of 75% for the full sample of 677 cluster candidates. Under the assumption that there are no false associations between the identified optical/NIR galaxy overdensities and SPT detections (we estimate that $<4\%$ of candidates will have such a false association—see discussion in Section 4.2 in S12), the measured cumulative purity of the sample is in excellent agreement with

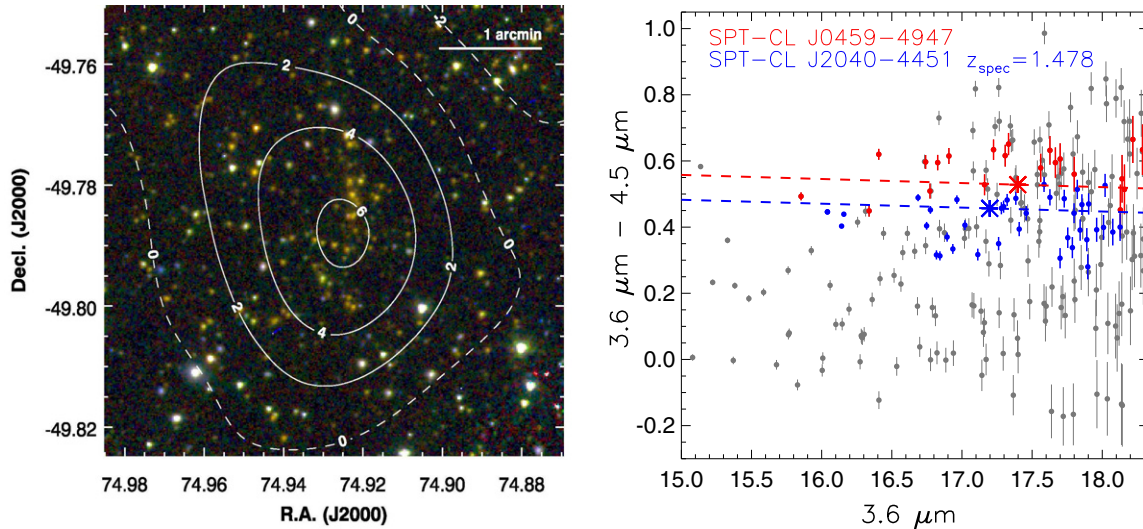


Figure 4. Left: SPT-CL J0459–4947, one of three SPT clusters with an estimated redshift of $z > 1.5$ (*rgb*: *Spitzer*/IRAC 4.5, 3.6 μm , Magellan/IMACS z band; over-plotted are the contours of the SZ detection). Right: *Spitzer* color–magnitude diagram (with magnitudes relative to Vega): plotted in gray are all galaxies in the *Spitzer* field; over-plotted in red are the galaxies identified with the SZ detection. The galaxies of this massive system ($M_{500c} \sim 3 \times 10^{14} M_{\odot} h_{70}^{-1}$) have significantly redder *Spitzer* colors than spectroscopically confirmed SPT-CL J2040–4451 at $z = 1.478$ (blue), supporting that its redshift is greater than $z = 1.5$. The color–magnitude relations of the best-fit model redshifts are over-plotted as dashed lines and the locations of model L^* galaxies are indicated via “*.” The best-fit redshift is $z = 1.7 \pm 0.2$, but the model is poorly calibrated at such high redshifts.

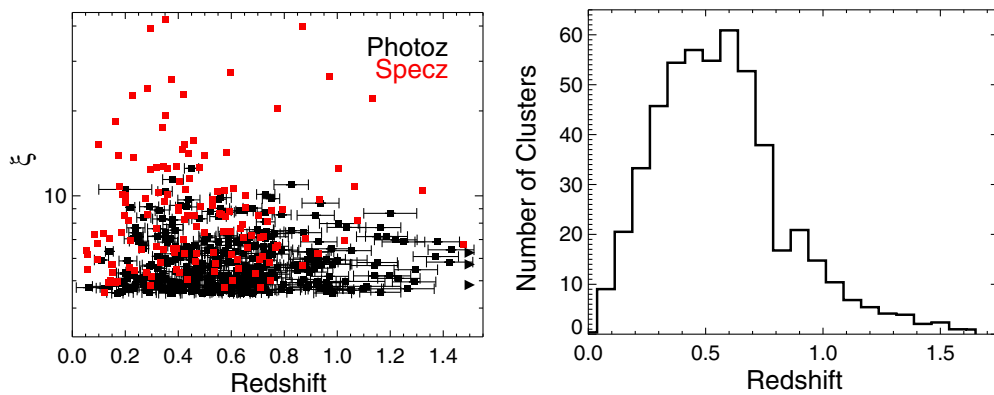


Figure 5. Left: plot of maximum detection significance, ξ , vs. redshift for the confirmed SPT-SZ cluster sample; black points correspond to systems with photometrically estimated redshifts while red points represent spectroscopically confirmed clusters. We report lower limits for the redshifts of the three highest-redshift systems (see Section 5.3). Right: the redshift distribution of the confirmed cluster sample; the median redshift of the sample is $z = 0.55$. The histogram does not have integer values as clusters with photometric redshift uncertainties were distributed amongst the appropriate bins.

simulations: the purity is $\geq 95\%$ at $\xi \geq 5$ and $\geq 76\%$ for the entire sample at $\xi \geq 4.5$. Here we quote the purity as a lower limit, as unconfirmed candidates may be clusters at redshifts too high to be confirmed with our follow-up imaging (and some lower-significance candidates have not yet been imaged).

The SPT-SZ cluster sample contains massive galaxy clusters over a wide redshift range. The median mass of the sample is $M_{500c} \sim 3.5 \times 10^{14} M_{\odot} h_{70}^{-1}$, and the median redshift is $z = 0.55$. The sample extends from $0.047 \leq z \lesssim 1.7$, and the mass threshold of the catalog is nearly independent of redshift (see Figure 6). This implies that the catalog reported here contains all of the most massive clusters in the $\sim 1/16$ th of the sky imaged by the SPT.

In Figure 7, we show the estimated selection function of the cluster sample in three redshift bins. Because ξ is the selection variable, the selection function can simply be written as the Heaviside step function $\Theta(\xi - 4.5)$. Given the ξ – M relation discussed in Section 6.1, we transform this function to mass space, where it represents the probability of a cluster of a

given mass to be included in the SPT-SZ sample. Note that at $M_{500c} > 7 \times 10^{14} h_{70}^{-1} M_{\odot}$ and $z > 0.25$, the SPT-SZ cluster catalog presented in this work is highly complete, meaning that nearly every such cluster in the surveyed area is present in the catalog. In Figure 6, we compare the mass and redshift distribution of the SPT sample to those from other large cluster catalogs selected via their ICM observables: namely the clusters detected in the all-sky *ROSAT* survey (Piffaretti et al. 2011), which includes the NORAS (Böhringer et al. 2000), REFLEX (Böhringer et al. 2004), and MACS (Ebeling et al. 2001) cluster catalogs; the 861 confirmed clusters from the all-sky *Planck* survey (Planck Collaboration et al. 2014a); and the 91 clusters that comprise the ACT cluster sample (Marriage et al. 2011; Hasselfield et al. 2013).

The mass threshold of the SPT sample declines slightly as a function of redshift owing to a combination of effects. At low redshifts ($z < 0.3$), increased power at large angular scales from primary CMB fluctuations and atmospheric noise raises the mass threshold for a fixed ξ cutoff (see, e.g., Vanderlinde

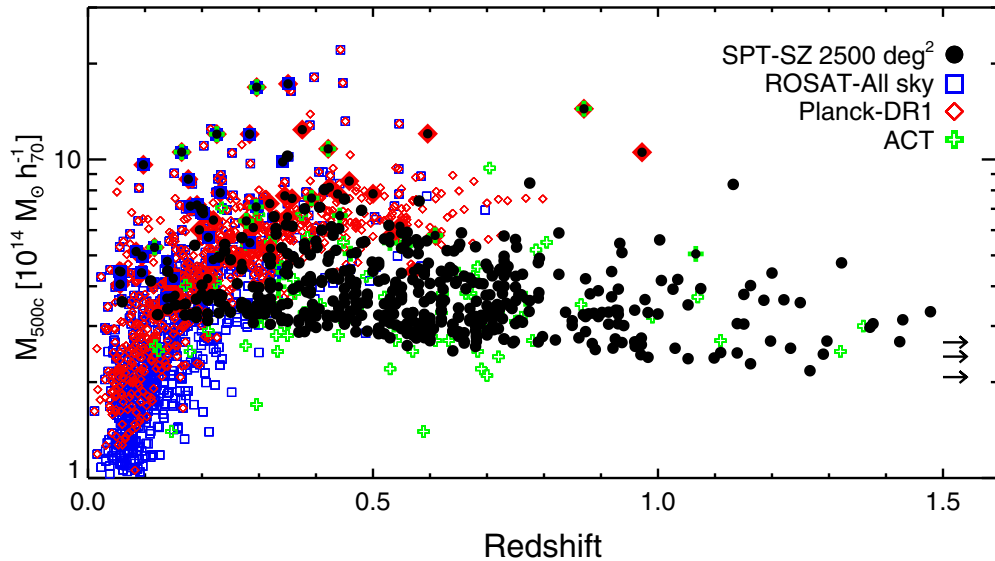


Figure 6. Comparison of the 2500 deg² SPT-SZ cluster catalog to other X-ray and SZ-selected cluster samples. Here we plot the estimated mass vs. redshift for the 516 optically confirmed clusters from the SPT catalog, 91 clusters from the ACT survey (Marriage et al. 2011; Hasselfield et al. 2013), 809 SZ-selected clusters from the *Planck* survey (Planck Collaboration et al. 2014a), and 740 X-ray clusters selected from the *ROSAT* all-sky survey (Piffaretti et al. 2011) with $M_{500c} \geq 1 \times 10^{14} h_{70}^{-1} M_{\odot}$. We mark 68% confidence lower limits for the redshifts of the three high-redshift SPT systems for which the *Spitzer* redshift model is poorly constrained (right arrows). We plot clusters in common between SPT and the other data sets (see, e.g., Table 5) at the SPT mass and redshift and, for common clusters in the other data sets, at the mass and redshift of the data set in which the cluster was first reported. While the SPT data provides a nearly mass-limited sample, the cluster samples selected from *ROSAT* and *Planck* data are redshift-dependent owing to cosmological dimming of X-ray emission and the dilution of the SZ signal by the large *Planck* beams, respectively.

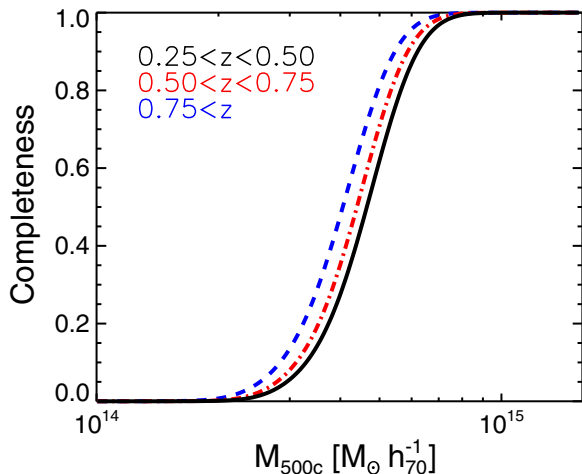


Figure 7. Completeness fraction as a function of mass for the SPT cluster sample in three different redshift bins: $0.25 < z < 0.50$ (solid black), $0.5 < z < 0.75$ (dashed blue), $z > 0.75$ (dotted red). The SPT sample is expected to be nearly 100% complete for $M_{500c} > 7 \times 10^{14} h_{70}^{-1} M_{\odot}$ at $z > 0.25$. Masses are calculated for a fiducial flat Λ CDM cosmology with $\sigma_8 = 0.80$, $\Omega_b = 0.046$, $\Omega_c = 0.254$, $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\tau = 0.089$, and $n_s(0.002) = 0.972$. Adopting the best-fit *Planck* cosmology (Planck Collaboration et al. 2014b) shifts the mass thresholds up $\sim 17\%$.

et al. 2010), while at higher redshifts the detectability of clusters is enhanced owing to increased temperatures for clusters of fixed mass. However, both of these trends are shallow, and the nearly redshift-independent selection function of the SPT catalog stands in contrast to the strong redshift dependence in X-ray catalogs and the *Planck* sample. The mass threshold for X-ray catalogs is redshift-dependent owing to cosmological dimming of the X-ray emission, while the redshift dependence of the *Planck* sample is driven by the dilution of the small

angular-scale signal of high-redshift clusters by the large *Planck* beam ($7'$ at 143 GHz).

We search the literature for counterparts to SPT candidates. We query the SIMBAD⁵⁶ and NED⁵⁷ databases as well as the union catalog of SZ sources detected by *Planck* (Planck Collaboration et al. 2014a) for counterparts. For confirmed clusters with $z \leq 0.3$ we utilize a $5'$ association radius; otherwise we match candidates within a $2'$ radius. All matches are listed in Table 5; we discuss potential false associations in the footnotes of this table. Additionally, we associate the brightest cluster galaxies in two clusters (SPT-CL J0249–5658 and SPT-CL J2254–5805) with spectroscopic galaxies from the 2dF Galaxy Redshift Survey (Colless et al. 2003) and the 6dF Galaxy Survey (Jones et al. 2009), respectively. In total, 115 of the SPT candidates are found to have counterparts in the literature (14 of these clusters were first discovered in SPT data). We report the new discovery of 251 clusters here, increasing the number of clusters first discovered in SPT data to 415. We highlight particularly noteworthy systems below, and a subset of the SPT cluster catalog is shown in Figure 8.

6.1. Cluster Mass Estimates

We provide estimated masses for all confirmed clusters in Table 4. These estimates, determined from each cluster's ξ and redshift, are based upon the methodology presented in Benson et al. (2013) and R13 but are reported here for a fixed flat Λ CDM cosmology—with $\sigma_8 = 0.80$, $\Omega_b = 0.046$, $\Omega_m = 0.30$, $h = 0.70$, $\tau = 0.089$, and $n_s(0.002) = 0.972$ —and a fixed ξ -mass scaling relation. In this section we provide a brief overview of the method; readers are referred to the earlier publications for additional details.

⁵⁶ <http://simbad.u-strasbg.fr/simbad>

⁵⁷ <http://nedwww.ipac.caltech.edu/>

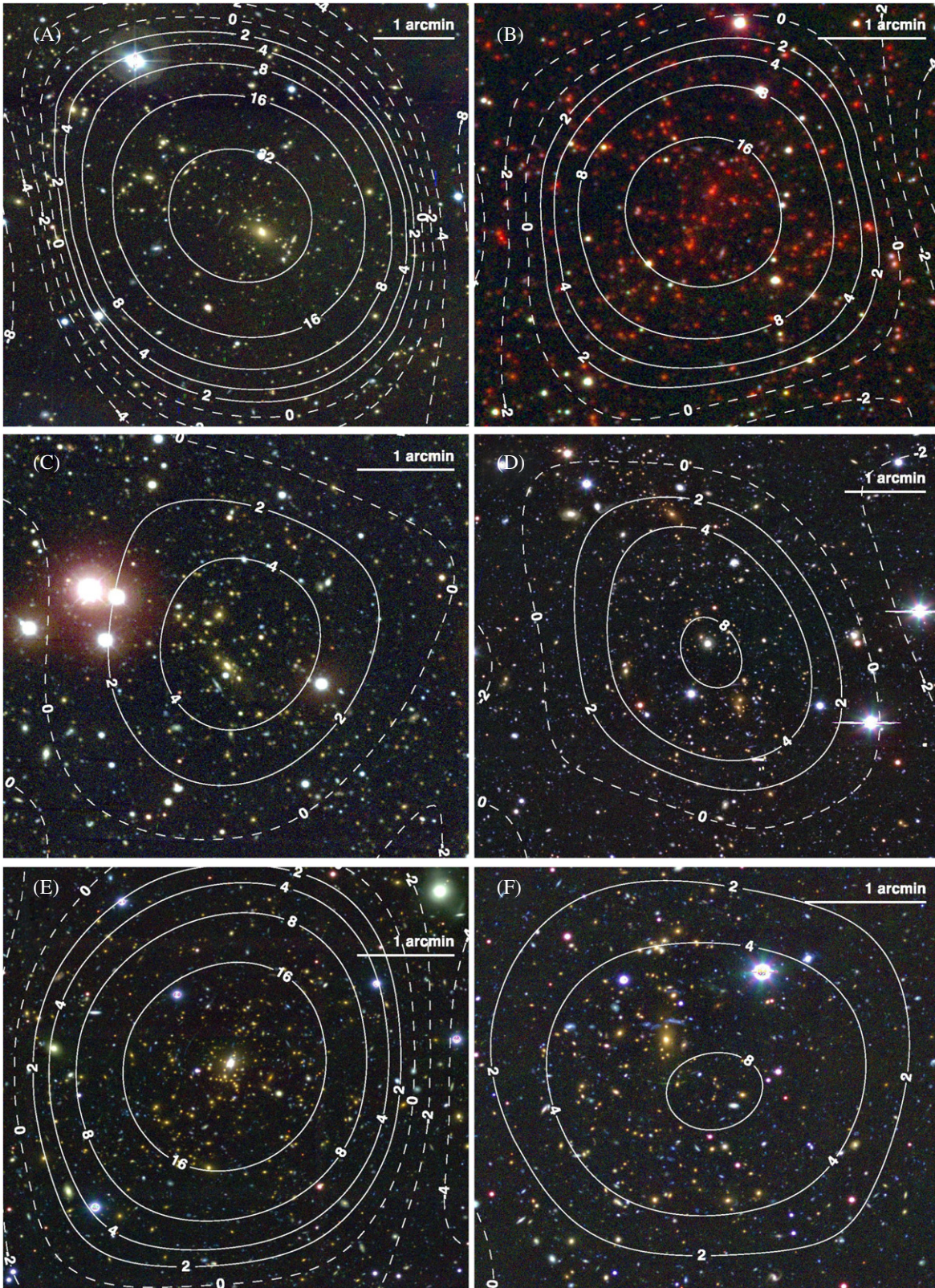


Figure 8. Sample of clusters from the 2500 deg² SPT-SZ cluster catalog. For each cluster we display an optical/NIR *rgb* image with the SZ detection contours over-plotted; see Section 6.4 for more details on particularly notable systems. (A) SPT-CL J2248–4431 (ACO S1063; $\xi = 42.4$, $z = 0.351$). This cluster is the most significant detection in the SPT sample (MPG/ESO WFI *IRV*-band image). (B) SPT-CL J2106–5844 ($\xi = 22.2$, $z = 1.132$)—also shown in SPT millimeter-wave data in Figure 1—is the most massive known cluster at $z > 1$. (*Spitzer*/IRAC 3.6 μm , Magellan/FourStar *J*-band, Magellan/IMACS *i*-band image) (C) SPT-CL J0410–6343 ($\xi = 5.6$, $z = 0.52$) is a “typical” SPT cluster at approximately the median redshift and ξ of the confirmed cluster sample. (Blanco/MOSAIC-II *gri*-band image). (D) SPT-CL J0307–6225 ($\xi = 8.5$, $z = 0.581$) is undergoing a major merger. As SZ selection is not greatly influenced by mergers or complicated astrophysics at the cores of clusters (e.g., Motl et al. 2005; Fabjan et al. 2011), the SPT sample is representative of the entire population of massive clusters (Magellan/Megacam *gri*-band image). (E) SPT-CL J2344–4243 (the “Phoenix Cluster”; $\xi = 27.4$, $z = 0.596$) is the most X-ray luminous cluster known. We confirm this cluster as a strong lens using newly acquired Megacam imaging (Magellan/Megacam *gri*-band image). (F) SPT-CL J0307–5042 ($\xi = 8.4$, $z = 0.55$) is one of many strong-lensing clusters in the SPT sample (Magellan/Megacam *gri*-band image).

To estimate each cluster’s mass, we compute the posterior probability density function:

$$P(M|\xi) \propto \frac{dN}{dMdz} \Big|_z P(\xi|M), \quad (11)$$

where $dN/dMdz$ is our assumed cluster mass function (Tinker et al. 2008), and $P(\xi|M)$ denotes the ξ -mass scaling relation. We assume an observable-mass scaling relation of the form

$$\zeta = A_{SZ} \left(\frac{M_{500c}}{3 \times 10^{14} M_{\odot} h^{-1}} \right)^{B_{SZ}} \left(\frac{H(z)}{H(0.6)} \right)^{C_{SZ}}, \quad (12)$$

parameterized by the normalization A_{SZ} (corrected field-by-field for the different noise levels in each field, see Table 1 and R13), the slope B_{SZ} , the redshift evolution C_{SZ} (where $H(z)$ is the Hubble parameter), and a log-normal scatter on ζ , D_{SZ} , where ζ is the “unbiased SPT-SZ significance”⁵⁸

$$\zeta = \sqrt{\langle \xi \rangle^2 - 3} \quad (13)$$

for $\zeta > 2$.

We fix the scaling relation parameters to $A_{SZ} = 4.14$, $B_{SZ} = 1.44$, $C_{SZ} = 0.59$, and $D_{SZ} = 0.22$. These values are the best-fit weighted averages as determined from a Monte Carlo Markov chain analysis of the R13 data set assuming a fixed scatter of 0.22 and the above canonical cosmology.⁵⁹ We caution that the masses for low-redshift clusters ($z < 0.25$) may be underestimated, and for low-significance clusters ($4.5 < \xi < 5$) the mass estimates should be considered only approximate. At low redshift the SZ signal becomes CMB-confused and therefore fails to obey the power-law form of the scaling relation (Vanderlinde et al. 2010). There is a more subtle complication for low-significance clusters. When we compute $P(M|\xi)$ in Equation (11), the theoretical halo mass function is used as the Bayesian prior. This choice implies a one-to-one mapping between halos and ξ values. However, this assumption breaks down for lower-mass halos, the total number of which approaches the number of independent resolution elements in the filtered SPT maps. Consequently, the contribution of these lower mass systems to $P(M|\xi)$ is overestimated. As we have already confirmed the existence of a massive system by requiring a significant red-sequence galaxy overdensity at the cluster location, we place a prior of $M_{500c} > 1 \times 10^{14} M_{\odot} h^{-1}$ when computing mass estimates. Decreasing this prior to $M_{500c} > 5 \times 10^{13} M_{\odot} h^{-1}$ typically shifts the mass of the lowest significance systems by less than 0.2σ .

Because we have used a fixed cosmology and scaling relation, the uncertainty reported on each cluster’s mass estimate only includes the contributions from measurement noise and the intrinsic scatter in the mass-observable relation. We also expect a comparable level of systematic uncertainty due to uncertainties in cosmology and scaling-relation parameters. This systematic uncertainty will be largely correlated between clusters, and is dominated by the uncertainty in the ξ -mass relation and our choice of external cosmological data sets. Here we have intrinsically linked the cluster mass estimates to our chosen cosmology by requiring the measured R13 cluster abundance to be consistent with this model. Assuming different cosmologies can shift the cluster mass scale at a level comparable to the

statistical uncertainty on the mass estimates. For example, adopting the best-fit Λ CDM model determined in R13 lowers the mass estimates by $\sim 8\%$ on average, whereas assuming parameter values consistent with the CMB data from WMAP9 ($\sigma_8 = 0.83$, $\Omega_m = 0.28$, $h = 0.70$; Hinshaw et al. 2013) or *Planck* ($\sigma_8 = 0.84$, $\Omega_m = 0.32$, $h = 0.67$; Planck Collaboration et al. 2014b) typically increases the mass estimates by $\sim 4\%$ and 17% , respectively.

We would prefer to observationally calibrate the cluster scaling relations and to independently constrain cosmological parameters using clusters. To achieve this goal, the SPT collaboration has undertaken a multi-wavelength campaign to obtain X-ray, galaxy velocity dispersion, and weak lensing measurements for ~ 50 – 100 clusters per technique. Early results from this work have been presented in other SPT publications, including Benson et al. (2013), R13, and Bocquet et al. (2014). In dH15 we will present cosmological constraints from the full 2500 deg^2 SPT-SZ cluster sample. This analysis will combine SZ and X-ray observations (we have obtained *Chandra* data for a significant subset of clusters, see Table 4 and Section 6.4.1) to both constrain cosmological parameters and better quantify the systematic uncertainties in mass estimates for the cluster sample.

6.2. Cluster Candidates in the Point-source-masked Regions

The point-source veto discussed in Section 3 rejects any cluster detections within $8'$ of an emissive source detected above 5σ at 150 GHz in SPT-SZ data. A total area of 163 square degrees was excluded from cluster finding for this reason. While such a conservative approach is appropriate when the goal is a cluster catalog with a clean selection function and a mass-observable relation with minimal outliers, it will almost certainly result in some massive clusters being excluded from the catalog. Assuming no spatial correlation between emissive sources and clusters, we would expect roughly 25 missed clusters above $\xi = 5$.

As in R13, we re-ran the cluster-finding algorithm only masking sources above $S_{150\text{GHz}} = 100 \text{ mJy}$ (as opposed to the normal threshold of $\sim 6 \text{ mJy}$). Each detection with $\xi \geq 5$ and with no counterpart in the original, conservatively masked catalog was visually inspected. The vast majority of these detections were rejected as obvious point-source-related artifacts, but some were clearly significant SZ decrements only minimally affected by the nearby source. These objects are listed in Table 3. We find 19 objects above $\xi = 5$, within roughly 1σ of the naive expectation, assuming purely Poisson statistics. All six candidates that were identified in the analogous procedure in R13 were also identified in this search.⁶⁰

As in R13, these auxiliary candidates were not included in our optical/IR follow-up campaign, and they are not included in the cosmological analysis or in the total number of candidates quoted in the rest of the text. We have searched the literature for counterparts to these clusters in other catalogs, and any literature counterparts are listed in Table 3. We also searched the RASS bright- and faint-source catalogs (Voges et al. 1999, 2000) for X-ray counterparts, and we show in Table 3 whether we find a RASS counterpart within either a $2'$ or $5'$ radius.

⁵⁸ See Appendix B in Vanderlinde et al. (2010)

⁵⁹ This scatter was chosen to be consistent with previous constraints from X-ray measurements (Benson et al. 2013; R13).

⁶⁰ There is a systematic ~ 0.75 offset between the positions of the R13 clusters found in source-masked areas and their counterparts in Table 3. This is due to a small error in the position calculation in R13. That error was only present in the source-masked cluster positions; as noted in Section 6.3, the positions in the main catalog here and the main catalog in R13 are in excellent agreement.

Table 3
Cluster Candidates above $\xi = 5$ in the Source-masked Area

SPT ID	R.A.	Decl.	ξ	θ_c	X-Ray Counterpart?		Literature Name
					within 2'	within 5'	
SPT-CL J0003–5253	0.8237	–52.8970	5.37	2.75	N	N	...
SPT-CL J0115–5959	18.8096	–59.9887	5.91	0.25	Y	Y	...
SPT-CL J0205–4125	31.3274	–41.4224	7.05	1.50	N	N	...
SPT-CL J0222–5335	35.6926	–53.5970	5.16	0.25	N	N	...
SPT-CL J0245–5302	41.3805	–53.0359	15.95	0.50	Y	Y	ACO S0295 ^a
SPT-CL J0321–4515	50.3200	–45.2556	5.88	1.50	N	N	...
SPT-CL J0334–6008	53.6899	–60.1497	7.11	1.25	N	Y	...
SPT-CL J0336–4037	54.0689	–40.6314	9.94	0.75	Y	Y	ACO 3140 ^a
SPT-CL J0434–5727	68.6315	–57.4523	5.06	0.75	Y	Y	...
SPT-CL J0440–4744	70.2431	–47.7370	5.18	1.25	N	Y	...
SPT-CL J0442–5905	70.6496	–59.0929	6.27	0.25	N	N	...
SPT-CL J0616–5227	94.1430	–52.4546	7.81	0.75	Y	Y	ACT-CL J0616–5227 ^b
SPT-CL J2104–4120	316.0754	–41.3475	8.97	2.25	Y	Y	ACO 3739 ^a
SPT-CL J2142–6419	325.7063	–64.3224	9.44	0.25	N	N	...
SPT-CL J2154–5952	328.6973	–59.8821	7.10	0.50	N	N	...
SPT-CL J2154–5936	328.7003	–59.6068	6.31	0.50	Y	Y	...
SPT-CL J2246–5244	341.5844	–52.7430	5.67	3.00	Y	Y	ACO 3911 ^a
SPT-CL J2300–5100	345.1034	–51.0126	5.22	2.50	N	N	...
SPT-CL J2347–6246	356.8065	–62.7693	6.63	0.25	N	N	ACO 4036 ^a

Notes. Cluster candidates identified in a version of the analysis in which only the very brightest (>100 mJy) point sources are masked (see text for details). Only candidates from the area masked in the standard analysis are listed here. These candidates are not included in cosmological analyses or in the candidate numbers quoted in the text. X-ray counterparts are searched for in the RASS bright- and faint-source catalogs. Literature name and reference are given for the first known identification of the cluster.

^a Abell et al. (1989).

^b Marriage et al. (2011).

6.3. Comparison to Previous SPT Catalogs

The SPT collaboration has published three previous samples of galaxy clusters (Vanderlinde et al. 2010; Williamson et al. 2011, R13). The catalog we present here encompasses all of the data used in these previous analyses, so it is a potentially useful cross-check of our new analysis to compare to these earlier samples. The R13 sample included all the clusters in Vanderlinde et al. (2010), using the exact same data and analysis, so we do not perform a separate comparison to Vanderlinde et al. (2010) here. The catalog published in R13 was constructed using an analysis nearly identical to that used here; some small differences were pointed out in Section 2.2 and we describe the remaining differences here.

First, there is a small difference in the area in each field’s map over which clusters are extracted. This area is generally defined in SPT analyses as the set of pixels with total weight above a given fraction of the median weight in the map. To ensure full coverage of the 2500 deg² region (i.e., no gaps between fields), we use a slightly lower threshold in this work (70% of the median weight, as opposed to 80% in R13). This results in very slight differences in noise properties and, hence, ξ values for extracted cluster candidates.

For the three fields observed in 2009 (RA21HDEC–60, RA3H30DEC–60, and RA21HDEC–50), we use the same data as in R13, with the only significant analysis difference being the different field border definitions. Therefore, we expect the list of cluster candidates in those fields to be very similar to the corresponding list in R13. Indeed, above a signal-to-noise threshold of $\xi = 5.5$ in these fields, the cluster lists are identical between this work and R13, and the ratio of ξ values for these clusters between the two analyses is 1.00 ± 0.04 . Two clusters between $\xi = 5$ and $\xi = 5.5$ from R13 (SPT-CL J0411–5751 and SPT-CL J2104–5224) are not included in the catalog

presented here. This absence is due to an update in the point-source lists between the two analyses; both of the $\xi > 5$ R13 clusters missing from the current catalog are within 4’ of newly identified point sources. There are two clusters in these fields at $5 \leq \xi \leq 5.5$ in the current list that are not in the R13 sample (SPT-CL J2158–5451 and SPT-CL J2143–5509); these are a result of the redefined field boundaries.

For the two fields originally observed in 2008 and re-observed in 2010 or 2011 (RA5H30DEC–55 and RA23H30DEC–55), the analysis in R13 used only the 2008 data, whereas we use both years’ data here. We thus expect the cluster list from R13 in these fields to be a subset of the list from these fields in this work, and we expect the average ξ for those clusters to be higher in this work, by a factor related to the extra 150 GHz depth and added 95 GHz information. As expected, all clusters above $\xi = 6$ from R13 in these fields are also present in the sample from this work, and the mean ξ ratio between this work and R13 is 1.28 ± 0.10 for these objects. We note that this ratio should agree with the “field scaling factors” used in dH15 to rescale the normalization parameter in the ξ -mass relation. The average of the two fields’ scaling factors in dH15 is 1.35; the ξ ratio we determine here is $\sim 5\%$ lower than this value but consistent within 1σ .

There is one cluster above $\xi = 5$ from the RA5H30DEC–55 and RA23H30DEC–55 fields that drops out of the R13 catalog when we add the new data. The cluster candidate SPT-CL J2343–5521 (detected at $\xi = 5.74$ in R13) is detected at $\xi < 4$ in the full two-year, two-band data. This trend of significance with added data is strong evidence that this is a false detection in the single-year, single-frequency data. This was already the preliminary conclusion in Vanderlinde et al. (2010) and R13, based on null results from optical and X-ray follow-up observations of this candidate.

We also compare the best-fit positions for the clusters detected in common between R13 and this work. The mean difference in

Table 4
Galaxy Clusters above $\xi = 4.5$ in 2500 deg² Observed by the SPT

SPT ID	R.A.	Decl.	Best		$Y_{SZ} \times 10^6$ (arcmin ²)	Redshift	M_{500c} ($10^{14} h_{70}^{-1} M_{\odot}$)	Imaging	Notes
	(J2000)	(J2000)	ξ	θ_c					
SPT-CL J0000–4356	0.0663	−43.9494	5.92	0.25	76 ± 21	1.00 ± 0.11	3.27 ± 0.71	2	
SPT-CL J0000–5748*	0.2499	−57.8064	8.49	0.50	82 ± 12	0.702 ⁺	4.56 ± 0.80	2	a
SPT-CL J0001–4842	0.2768	−48.7132	5.69	1.25	77 ± 12	0.30 ± 0.04	3.96 ± 0.87	5	
SPT-CL J0001–4024	0.3610	−40.4108	5.42	0.75	59 ± 12	0.85 ± 0.07	3.12 ± 0.75	2, 10	
SPT-CL J0001–6258	0.4029	−62.9808	4.69	1.50	51 ± 16	0.21 ± 0.02	3.27 ± 0.91	3	b, c
SPT-CL J0001–5440	0.4059	−54.6697	5.69	1.00	60 ± 12	0.73 ± 0.06	3.54 ± 0.79	1, 10	
SPT-CL J0002–5557	0.5138	−55.9621	5.20	0.25	54 ± 18	1.14 ± 0.08	2.47 ± 0.60	3, 10	
SPT-CL J0002–5224	0.6433	−52.4092	4.67	1.00	48 ± 12	>0.71	...	2	
SPT-CL J0003–4155	0.7842	−41.9307	4.75	0.25	58 ± 20	>0.76	...	3	
SPT-CL J0007–4706	1.7514	−47.1159	4.55	0.75	58 ± 13	0.51 ± 0.04	2.95 ± 0.85	5	b

Notes. Galaxy cluster candidates selected above a significance of 4.5 in the 2500 deg² SPT-SZ survey. Galaxy clusters marked by a “*” have X-ray data from the *Chandra* satellite that are used in the cosmological analysis presented in dH15. For each candidate we report the position, the highest detection significance in the filtered maps and the core radius corresponding to the detection (in arcmin; see Section 3), the integrated Y_{SZ} within a 0.75 aperture, the redshift for confirmed systems (except for the three highest-redshift systems; see Section 5.3) or redshift lower limit for unconfirmed systems (see Section 4), and a mass estimate for each confirmed cluster (Section 6.1). Spectroscopic redshifts are quoted without uncertainties, which are typically $\sim 0.1\%$ (see, e.g., Ruel et al. 2014). The spectroscopic redshifts derived from SPT follow-up observations are marked with a “+.”

^a Strong-lensing cluster.

^b Cluster masses at low significance are only approximate; see Section 6.1.

^c Cluster masses at low redshift ($z < 0.25$) may be underestimated; see Section 6.1.

^d Confirmed, high-redshift cluster (see Section 5.3).

¹ Bright star impedes confirmation.

² Possibly biased SZ-center owing to cluster’s proximity to the edge of the cluster map; additional cluster offset 1.8 at $z \sim 0.55$.

³ Foreground group at $z \sim 0.3$.

⁴ Foreground group to N at $z \sim 0.35$.

⁵ Foreground group at $z \sim 0.1$.

⁶ Group $z \sim 0.2$ centered 3’ away from SPT center.

⁷ Small group $z \sim 0.4$ offset 2’ to N

⁸ Low- z group $z \sim 0.1$.

⁹ Very complex region; optical group on NW $z \sim 0.4$ and another group on SW $z \sim 0.65$. Confirmed in S12 by one method at $z = 0.38 \pm 0.04$.

¹⁰ Bimodal redshift solution; additional solution $z = 0.29$.

¹¹ High stellar density.

¹² Foreground group at $z \sim 0.15$.

¹³ Bimodal redshift solution; additional solution $z = 0.41$.

¹⁴ Group on SW $z \sim 0.4$.

¹⁵ System $z \sim 0.3$ offset 2’ to S.

¹⁶ Possibly biased SZ-center owing to cluster’s proximity to the edge of the cluster map.

¹⁷ The mass is biased low by a factor of $\sim 1.5\times$ owing to contamination from a magnified high-redshift dusty star-forming galaxy; see Section 6.4.3 and discussion in Andersson et al. (2011).

¹⁸ Low- z group at $z \sim 0.15 > 3'$ from SPT center.

(This table is available in its entirety in machine-readable form.)

positions of clusters in the two catalogs is $0'.15 \pm 0'.59$ in right ascension and $-0'.04 \pm 0'.69$ in declination.

We also cross-check the cluster catalog presented in Williamson et al. (2011) with the catalog presented here. Williamson et al. (2011) used full-depth observations of 60% of the SPT-SZ survey area and “preview-depth” (roughly three times the nominal noise level) observations of the rest of the survey to compile a catalog of the most massive clusters over the full 2500 deg² region. All clusters detected in that work are also detected here. One cluster from Williamson et al. (2011)—SPT-CL J0245–5302, also known as ACO S0295—is near a bright radio source and is thus not included in the official catalog in Table 4; it is, however, detected in the alternative analysis in Section 6.2 that avoids only the very brightest point sources, and it is included in Table 3. The values of ξ in this work for the clusters from full-depth data in Williamson et al. (2011) are consistent with the ξ values reported in that work; for clusters found in preview-depth data in Williamson et al. (2011), the ξ values reported here have increased by roughly a factor of two over that work.

Finally, we note that there may be small differences in redshifts/redshift limits and the confirmation status of cluster candidates. These changes are driven by a mixture (in varying degree per candidate) of additional follow-up data, improved optical data processing and re-estimated redshifts (Section 5) calibrated using the enlarged SPT spectroscopic sample (Section 4.2.3).

6.4. Notable Clusters

In this section we highlight particularly notable clusters and subsets of clusters from the SPT-SZ cluster catalog.

6.4.1. The SPT-XVP Sample

Eighty of the most significant SPT clusters discovered in the first 2000 deg² at $z > 0.4$ have been observed by *Chandra* as part of a large X-ray Visionary Project (XVP; PI: Benson). As described in R13 and dH15, these observations, with ~ 2000 X-ray counts/cluster, play a critical role in constraining the SZ-mass scaling relation and greatly strengthen the

Table 5
Clusters with Matches in Other Catalogs

SPT ID	First ID, Ref.	All Catalogs with Match	z	Lit. z , Ref.
SPT-CL J0010–5112	ACO S0013, 1	1	0.17 ± 0.02	...
SPT-CL J0013–4621	[BM78] 210, 2	1, 2	0.18 ± 0.05	...
SPT-CL J0027–5015	Str 0025–505, 3	1, 3, 4, 5, 6	0.145	0.1448, [4]
SPT-CL J0036–4411	SpARCS J003645–441050, 7	7	0.869	0.869, [7]
SPT-CL J0041–4428	ACO S0067, 1	1	0.33 ± 0.02	...
SPT-CL J0051–4834	Str 0049–489, 3	1, 3	0.187	0.1873, [8]
SPT-CL J0102–4915	ACT-CL J0102–4915, 9	6, 9, 10	0.870	0.8701, [11]
SPT-CL J0108–4341	ACO 2873, 1	1	0.19 ± 0.02	...
SPT-CL J0118–5638	[BM78] 148, 2	1, 2	0.21 ± 0.04	...
SPT-CL J0124–5937	PSZ1 G295.19–56.99, 6	6	0.21 ± 0.03	0.22, [6]
SPT-CL J0129–6432	RXC J0129.4–6432, 12	12	0.326	0.3264, [12]
SPT-CL J0133–6434	PSZ1 G295.60–51.95, 6	6	0.29 ± 0.04	0.33, [6]
SPT-CL J0135–5904	PSZ1 G292.40–57.11, 6	6	0.49 ± 0.03	...
SPT-CL J0145–5301	Sersic 017/04, 13	1, 4, 5, 9, 13, 14	0.117	0.1168, [4]
SPT-CL J0145–6033	RBS 0238, 15	4, 6, 15	0.179	0.1795, [16]
SPT-CL J0214–4638	QW 0213–468, 17	17	0.27 ± 0.02	...
SPT-CL J0216–4816	ACO 2998, 1	1, 8, 18	0.171	0.1709, [8]
SPT-CL J0217–5245	RXC J0217.2–5244, 4	4, 9	0.343	0.3432, [4]
SPT-CL J0225–4155	ACO 3017, 1	1, 4, 8, 14, 19, 20	0.220	0.2195, [8]
SPT-CL J0232–4421	[DBG99] 27, 14	8, 10, 14, 15, 20	0.284	0.2836, [14]
SPT-CL J0232–5257	ACT-CL J0232–5257, 9	9	0.556	0.5559, [21]
SPT-CL J0235–5121	ACT-CLJ0235–5121, 9	6, 9	0.278	0.2777, [21]
SPT-CL J0236–4938	ACT-CLJ0237–4939, 9	9	0.334	0.4, [22]
SPT-CL J0243–4833	SPT-CL J0243–4833, 10	6, 10	0.500	0.53, [10]
SPT-CL J0254–5857	SPT-CL J0254–5856, 10	10, 20, 23, 24	0.438	0.438, [24]
SPT-CL J0256–4736	ACO 3072, 1	1	0.23 ± 0.03	...
SPT-CL J0304–4401	SPT-CL J0304–4401, 10	6, 10	0.458	0.52, [10]
SPT-CL J0304–4921	ACT-CL J0304–4921, 9	6, 9	0.392	0.3922, [21]
SPT-CL J0307–4123	PSZ1 G248.96–58.75, 6	6	0.67 ± 0.04	...
SPT-CL J0311–6354	PSZ1 G281.29–46.89, 6	6	0.25 ± 0.02	0.28, [6]
SPT-CL J0317–4849	ACO 3113, 1	1, 6, 8	0.164	0.1642, [25]
SPT-CL J0328–5541	ACO 3126, 1	1, 4, 14, 15, 18, 20, 23, 26	0.084	0.0844, [27]
SPT-CL J0330–5228	ACT-CLJ0330–5227, 9	9	0.442	0.44, [22]
SPT-CL J0346–5439	ACT-CL J0346–5438, 9	9	0.530	0.5297, [21]
SPT-CL J0348–4515	CIG 0346–4524, 28	17, 28	0.358	0.3251, [28]
SPT-CL J0404–6510	[BM78] 082, 2	1, 2, 6, 23	0.12 ± 0.03	0.11, [29]
SPT-CL J0411–4819	SPT-CL J0411–4819, 10	10, 20	0.422	0.42, [10]
SPT-CL J0411–6340	ACO 3230, 1	1, 15	0.14 ± 0.01	0.14, [29]
SPT-CL J0431–6126	Sersic 040/06, 13	1, 3, 4, 13, 14, 18, 20, 23, 26, 30, 31, 32	0.058	0.0577, [27]
SPT-CL J0438–5419	ACT-CL J0438–5419, 9	9, 10, 20	0.421	0.4214, [21]
SPT-CL J0458–5741	ACO 3298, 1	1, 23	0.19 ± 0.02	...
SPT-CL J0500–5116	ACO 3303, 1	1, 33	0.11 ± 0.03	0.14, [25]
SPT-CL J0504–4929	ACO 3311, 1	1	0.20 ± 0.03	...
SPT-CL J0505–6145	PSZ1 G271.28–36.12, 6	6	0.25 ± 0.03	...
SPT-CL J0509–5342	SPT-CL J0509–5342, 34	9, 10, 23, 34, 35, 36	0.461	0.4626, [37]
SPT-CL J0509–6118	PSZ1 G270.64–35.68, 6	6	0.35 ± 0.03	0.31, [6]
SPT-CL J0510–4519	ACO 3322, 1	1, 4, 15, 17, 20, 38	0.20	0.20, [27]
SPT-CL J0511–5154	SCSO J051145–515430, 39	23, 35, 39	0.645	0.645, [24]
SPT-CL J0512–5139	SCSO J051240–513941, 39	36, 39	0.60 ± 0.03	0.66, [39]
SPT-CL J0516–5430	ACO S0520, 1	1, 4, 9, 20, 23, 34, 35, 36, 39	0.295	0.2952, [4]
SPT-CL J0519–4248	RCS J051919–4247.8, 40	40	0.59 ± 0.03	0.6, [40]
SPT-CL J0521–5104	SCSO J052113–510418, 39	21, 23, 35, 36, 39	0.675	0.6755, [21]
SPT-CL J0522–4818 ^a	DLSCL J0522.2–4820, 41	1, 33, 41	0.296	0.296, [42]
SPT-CL J0522–5026	SCSO J052200–502700, 39	36, 39	0.52 ± 0.03	0.5, [39]
SPT-CL J0525–4715	ACO 3343, 1	1, 4, 20	0.191	0.1913, [4]
SPT-CL J0528–5300	SPT-CL J0528–5300, 34	9, 23, 34, 35, 39	0.768	0.7648, [43]
SPT-CL J0535–4801	PSZ1 G254.58–32.16, 6	6	0.92 ± 0.07	...
SPT-CL J0542–4100	RDCS J0542–4100, 44	44, 45	0.642	0.63, [44]
SPT-CL J0546–5345	SPT-CL J0547–5346, 34	9, 23, 34, 35	1.066	1.067, [46]
SPT-CL J0549–6205	SPT-CL J0549–6204, 10	10, 20	0.376	0.32, [10]
SPT-CL J0550–5019	LCS-CL J055019–5019.6, 36	36	0.65 ± 0.03	0.66, [36]
SPT-CL J0555–6406	SPT-CL J0555–6405, 10	6, 10	0.345	0.42, [10]
SPT-CL J0559–5249	SPT-CL J0559–5249, 35	6, 9, 23, 35	0.609	0.6112, [43]
SPT-CL J0603–4714	PSZ1 G254.54–27.32, 6	6	0.32 ± 0.04	0.27, [6]
SPT-CL J0611–5938	PLCKESZ G268.5–28.1, 47	47	0.39 ± 0.03	0.47, [47]
SPT-CL J0615–5746	SPT-CL J0615–5746, 10	6, 10, 48	0.972	0.972, [10]
SPT-CL J0628–4143	ACO 3396, 1	1, 4, 6, 10, 14	0.176	0.1759, [14]
SPT-CL J0637–4829	ACO 3399, 1	1, 4, 20	0.203	0.2026, [4]
SPT-CL J0638–5358	ACO S0592, 1	1, 4, 9, 10, 14, 20	0.226	0.2216, [14]
SPT-CL J0641–5001	ACO 3403, 1	1, 33	0.123	0.1226, [49]

Table 5
(Continued)

SPT ID	First ID, Ref.	All Catalogs with Match	z	Lit. z , Ref.
SPT-CL J0645–5413	ACO 3404, 1	1, 4, 9, 10, 14, 20, 33	0.164	0.167, [14]
SPT-CL J0651–4037	PSZ1 G250.29–17.29, 6	6	0.27 ± 0.02	0.27, [6]
SPT-CL J0658–5556	1E 0657–55.8 (Bullet), 50	4, 9, 10, 14, 20, 50	0.296	0.296, [50]
SPT-CL J2011–5725	RXC J2011.3–5725, 4	4, 23	0.279	0.2786, [4]
SPT-CL J2012–4130	[BM78] 236, 2	1, 2, 4, 6	0.150	0.1496, [4]
SPT-CL J2012–5649	Str 2008–569, 3	1, 2, 3, 4, 14, 15, 17, 20, 23, 31, 33	0.055	0.0556, [27]
SPT-CL J2020–4646	ACO 3673, 1	1, 23	0.19 ± 0.02	0.19, [23]
SPT-CL J2021–5257	Ser 138–5, 13	1, 3, 4, 13, 23	0.138	0.1383, [4]
SPT-CL J2023–5535	RXC J2023.4–5535, 4	4, 10, 20, 23	0.232	0.232, [4]
SPT-CL J2025–5117	[BM78] 188, 2	1, 2, 6, 23	0.22 ± 0.02	0.18, [23]
SPT-CL J2031–4037	RXC J2031.8–4037, 4	4, 10	0.342	0.3416, [4]
SPT-CL J2032–5627 ^b	[F81] 391, 51	1, 4, 6, 23, 33, 51	0.284	0.284, [24]
SPT-CL J2055–5456	[BM78] 164, 2	1, 2, 4, 23	0.139	0.139, [4]
SPT-CL J2059–5018	ACO S0912, 1	1, 23	0.32 ± 0.02	0.41, [23]
SPT-CL J2101–5542	ACO 3732, 1	1, 23	0.25 ± 0.02	0.2, [23]
SPT-CL J2121–6335	ACO S0937, 1	1, 23, 52	0.217	0.217, [52]
SPT-CL J2134–4238	[BM78] 253, 2	1, 2, 6, 31, 51	0.196	0.1955, [27]
SPT-CL J2138–6008	SPT-CL J2138–6007, 23	6, 23	0.319	0.319, [24]
SPT-CL J2148–6116	SPT-CL J2148–6116, 23	6, 23	0.571	0.571, [24]
SPT-CL J2201–5956	[F81] 408, 51	1, 4, 10, 14, 15, 18, 20, 23, 26, 51	0.097	0.0972, [27]
SPT-CL J2211–4833	[BM78] 202, 2	1, 2	0.24 ± 0.03	0.0576, [53]
SPT-CL J2217–6509	Sersic 150/06, 13	2, 3, 4, 13, 14, 17, 20	0.095	0.0951, [14]
SPT-CL J2223–5015	APMCC 769, 18	18	0.24 ± 0.04	...
SPT-CL J2223–5227	ACO 3872, 1	1	0.27 ± 0.03	...
SPT-CL J2241–4001	EIS J2241–4001, 54	54, 55	0.64 ± 0.03	1, [55]
SPT-CL J2241–4236	[BM78] 259, 2	1, 2, 8, 31	0.20 ± 0.03	...
SPT-CL J2248–4431	ACO S1063, 1	1, 4, 8, 10, 14, 15, 20	0.351	0.3475, [4]
SPT-CL J2249–6426	[BM78] 108, 2	1, 2, 4, 14, 18, 20, 26, 33	0.094	0.094, [56]
SPT-CL J2254–4620	ACO 3937, 1	1	0.27 ± 0.03	...
SPT-CL J2254–5805	[R76] 89, 57	3, 14, 57	0.153	...
SPT-CL J2254–6314	AM 2250–633, 31	4, 14, 31	0.211	0.2112, [14]
SPT-CL J2259–5617	[BM78] 165, 2	1, 2, 6, 23, 31, 35	0.153	0.17, [23]
SPT-CL J2300–5331	ACO S1079, 1	1, 23, 35	0.262	0.262, [24]
SPT-CL J2301–4023	[LP96] Cl2259–4040, 58	58	0.73 ± 0.04	...
SPT-CL J2306–6505	PSZ1 G319.20–48.61, 6	6	0.530	0.69, [6]
SPT-CL J2313–4243	Sersic 159/03, 13	1, 2, 3, 4, 6, 8, 13, 14, 15, 17, 38, 57	0.056	0.0564, [4]
SPT-CL J2316–5453	SCSO J231651–545356, 39	36, 39, 59	0.39 ± 0.05	0.44, [39]
SPT-CL J2325–4111 ^c	ACO S1121, 1	1, 6, 10	0.358	0.358, [24]
SPT-CL J2327–5137	LCS-CL J232708–5137.5, 36	36	0.32 ± 0.02	0.37, [36]
SPT-CL J2332–5358	SCSO J233227–535827, 39	23, 35, 36, 39	0.402	0.402, [24]
SPT-CL J2335–4544	PSZ1 G337.11–66.02, 6	6	0.547	0.5, [6]
SPT-CL J2344–4243	SPT-CL J2344–4243, 10	6, 10	0.596	0.596, [60]
SPT-CL J2350–5301	SCSO J235055–530124, 39	39	0.54 ± 0.03	0.46, [39]
SPT-CL J2351–5452	SCSO J235055–530124, 39	23, 36, 39	0.384	0.3838, [61]
SPT-CL J2354–5633	SCSO J235454–563311, 39	39	0.53 ± 0.03	0.51, [39]

Notes. Cluster candidates coincident with galaxy clusters identified in other catalogs. We define a match if a candidate is within $5'$ ($2'$) of an identified cluster for clusters at $z \leq 0.3$ ($z > 0.3$ or unconfirmed). For each match, we report the name under which the cluster was first reported and all catalogs which include the cluster. We also quote the cluster redshift used in this work—either the photometric redshift or a spectroscopic redshift obtained from follow-up observations or the literature. We include error bars for red sequence redshifts but not spectroscopic redshifts. In the last column, we quote a redshift from the literature if available. Error bars are not reported for literature redshifts; two (three or four) significant digits are used if the literature redshift is photometric (spectroscopic), with the exception of SPT-CL J0510–4519 for which the literature spectroscopic redshift is reported to two significant figures.

^a ACO 3338 is in the foreground.

^b [F81] 391/ACO 3685 is in the foreground.

^c ACO S1121 is in the foreground.

References. (1) ACO catalog, Abell et al. 1989; (2) Braid & MacGillivray 1978; (3) Stromlo catalog Duss & Newell 1977; (4) REFLEX catalog, Böhringer et al. 2004; (5) Panko & Flin 2006; (6) Planck Cluster catalog, Planck Collaboration et al. 2014a; (7) GCLASS catalog, Muzzin et al. 2012; (8) Cruddace et al. 2002; (9) ACT-CL Catalog, Marriage et al. 2011; (10) SPT-CL catalog, Williamson et al. 2011; (11) Menanteau et al. 2012; (12) REFLEX II Catalog, Chon & Böhringer 2012; (13) Sersic catalog, Sérsic 1974; (14) [DBG99] catalog, de Grandi et al. 1999; (15) RBS catalog, Schwoppe et al. 2000; (16) Guzzo et al. 2009; (17) [QW] catalog, Quintana & White 1990; (18) APMCC catalog, Dalton et al. 1997; (19) Edinburgh-Durham Cluster Catalog, Lumsden et al. 1992; (20) Planck Early SZ catalog, Planck Collaboration et al. 2011a; (21) Sifón et al. 2013; (22) Menanteau et al. 2010b; (23) SPT-CL catalog, Reichardt et al. 2013; (24) Ruel et al. 2014; (25) Coziol et al. 2009; (26) [DEM94] catalog, Dalton et al. 1994; (27) Struble & Rood 1999; (28) West & Frandsen 1981; (29) Ebeling et al. 1996; (30) Luggner 1978; (31) Arp & Madore 1996; (32) Edge et al. 1990; (33) Einasto et al. 1997; (34) SPT-CL catalog, Staniszewski et al. 2009; (35) SPT-CL catalog, Vanderlinde et al. 2010; (36) LCS Catalog, Bleem et al. 2015; (37) High et al. 2010; (38) Gioia et al. 1990; (39) Menanteau et al. 2010a; (40) Horesh et al. 2010; (41) Deep Lens Survey Cluster Catalog, Wittman et al. 2006; (42) Abate et al. 2009; (43) High et al. 2010; (44) Tozzi et al. 2003; (45) 400d Cluster Catalog, Burenin et al. 2007; (46) Brodwin et al. 2010; (47) Planck Collaboration et al. 2012; (48) Planck Collaboration et al. 2011b; (49) Coziol et al. 2009; (50) Tucker et al. 1998; (51) Feitsova 1981; (52) Webb et al. 2013; (53) Batuski et al. 1999; (54) Olsen et al. 1999; (55) Lobo et al. 2000; (56) Katgert et al. 1996; (57) Rose 1976; (58) Lidman & Peterson 1996; (59) XMM-BCS catalog, Šuhada et al. 2012; (60) McDonald et al. 2012; (61) Buckley-Geer et al. 2011.

cosmological constraining power of the SPT cluster sample. The broad redshift range of this data set has also enabled constraints on the redshift evolution of the X-ray properties of massive clusters, including measurements of their cooling properties (McDonald et al. 2013b) as well as the redshift evolution of the temperature, pressure and entropy profiles of clusters (McDonald et al. 2014a).

6.4.2. Massive Clusters at $z > 1$

The nearly redshift-independent selection of the SPT-SZ cluster sample has led to the discovery of a number of massive, high-redshift clusters. Thirty-seven clusters reported in this work have redshifts estimated at $z > 1$; three systems reported here for the first time: SPT-CL J0459–4947 ($\xi = 6.29$), SPT-CL J0446–4606 ($\xi = 5.71$) and SPT-CL J0334–4645 ($\xi = 4.83$) have redshifts estimated from *Spitzer* observations at $z > 1.5$. Several of the $z > 1$ systems have been the focus of more detailed study including: SPT-CL J0546–5345, the first $z > 1$ cluster detected by its SZ signature (Brodwin et al. 2010); SPT-CL J2106–5844, the most massive known cluster at $z > 1$ ($M_{500c} = 8.3 \times 10^{14} h_{70}^{-1} M_{\odot}$; Foley et al. 2011); SPT-CL J0205–5829 at $z = 1.32$ (Stalder et al. 2013), which features a red sequence whose bright galaxies are well-evolved by $z = 1.3$; and SPT-CL J2040–4451, which intriguingly shows signs of active star formation (Bayliss et al. 2014). At $z = 1.478$, SPT-CL J2040–4451 is the highest-redshift spectroscopically confirmed SPT cluster to date.

6.4.3. Strong Lensing Clusters

A number of SPT clusters can be identified from the literature and existing SPT follow-up observations as strong gravitational lenses. Previous SPT publications first identified SPT-CL J0509–5342, SPT-CL J0546–5345 (Staniszewski et al. 2009), SPT-CL J0540–5744, SPT-CL J2331–5051 (High et al. 2010), SPT-CL J2011–5228, and SPT-CL J2011–5725 (Song et al. 2012) as strong-lensing clusters using optical imaging data. SPT-CL J2332–5358 was identified as a lens by the presence of a multiply imaged, high-redshift ($z = 2.73$), DSFG (Greve et al. 2012; Aravena et al. 2013). The ACT team first reported the discovery of several clusters in the SPT sample (see Table 5) and identified 3 of these systems as strong lenses: SPT-CL J0304–4921, SPT-CL J0330–5228 (Menanteau et al. 2010a), and SPT-CL J0102–4915 (Menanteau et al. 2012; Zitrin et al. 2013). Other previously identified strong-lensing systems include SPT-CL J0658–5556 (1E0657-56/Bullet Cluster; Mehlert et al. 2001), SPT-CL 2031–4037 (RXJ J2031.8–4037; Christensen et al. 2012), SPT-CL J2248–4431 (ACO 1063S; Gómez et al. 2012), and SPT-CL J2351–5452 (SCSO J235055–530124; Menanteau et al. 2010b; Buckley-Geer et al. 2011).

In Table 4, we report 34 additional strong gravitational lenses. Optical images of two newly identified strong lenses (SPT-CL J2344–4243 and SPT-CL J2138–6008) are shown in Figure 8. While the majority of these lenses have been identified in ground-based imaging from the follow-up program described in Section 4, a subset has been identified via the presence of bright arcs in deeper, higher-quality imaging acquired as part of our weak-lensing mass calibration efforts (including 16 lensing clusters in data from the *Hubble Space Telescope*). We note that, given the heterogeneity in image quality in existing follow-up data, this list of strong lenses represents neither an exhaustive nor a uniformly selected sample of systems.

6.4.4. Notable Individual Systems

1. *SPT-CL J2248–4431*. First reported as ACO S1063, it is the most significant detection ($\xi = 42.4$, $z = 0.351$, $M_{500c} = 17.3 \times 10^{14} h_{70}^{-1} M_{\odot}$) in the SPT-SZ sample. The cluster is the most massive cluster in the SPT-SZ sample, and is the second most X-ray luminous cluster in the REFLEX X-ray catalog (Böhringer et al. 2004). It is scheduled for ultra-deep *Hubble Space Telescope* observations, as one of its Frontier Fields.⁶¹
2. *SPT-CL J0102–4915*. First reported in Menanteau et al. (2010a), this cluster is also known as “El Gordo” (Menanteau et al. 2012). Detected in the SPT-SZ survey at $\xi = 39.9$, this massive ($M_{500c} = 14.4 \times 10^{14} h_{70}^{-1} M_{\odot}$) merging cluster at $z = 0.870$ is the second most significant detection in the SPT-SZ sample. It has a very high X-ray temperature (14.5 keV), and an X-ray luminosity that makes it the second most X-ray luminous cluster in the SPT-SZ sample (Menanteau et al. 2012).
3. *SPT-CL J0658–5556*. This cluster is the well-known “Bullet” cluster (1ES 0657–558; $z = 0.296$, Tucker et al. (1998); Clowe et al. (2006)). Detected at $\xi = 39.0$, this cluster is the second most massive system ($M_{500c} = 16.9 \times 10^{14} h_{70}^{-1} M_{\odot}$) in the SPT-SZ cluster sample.
4. *SPT-CL J2344–4243*. This system, first reported in Williamson et al. (2011), is also known as the “Phoenix Cluster” ($\xi = 27.4$, $z = 0.596$, $M_{500c} = 12.0 \times 10^{14} h_{70}^{-1} M_{\odot}$). It is the most X-ray luminous cluster known in the universe. The properties of this system, including those of its central galaxy which exhibits an exceptionally high rate of star formation, are explored in detail in McDonald et al. (2012, 2013a, 2014b). We use newly acquired Magellan/Megacam imaging to identify this system as a strong lens (see Figure 8).
5. *SPT-CL J0615–5746*. This cluster ($\xi = 26.4$, $z = 0.972$, $M_{500c} = 10.5 \times 10^{14} h_{70}^{-1} M_{\odot}$) was first reported by SPT in Williamson et al. (2011), and also appears in the *Planck* cluster catalog (Planck Collaboration et al. 2011a). It has a measured X-ray luminosity equal to the Bullet cluster (Planck Collaboration et al. 2011b), but is at significantly higher redshift.
6. *SPT-CL J2106–5844*. As mentioned above, this cluster ($\xi = 22.2$, $z = 1.132$, $M_{500c} = 8.3 \times 10^{14} h_{70}^{-1} M_{\odot}$) is the most massive known cluster at $z > 1$. It has a measured X-ray luminosity nearly equal to SPT-CL J0615–5746, and is described in more detail in Foley et al. (2011).

7. CONCLUSIONS

This work has documented the construction and properties of a catalog of galaxy cluster candidates, selected via their SZ signature in the 2500 deg² SPT-SZ survey. Using a spatial-spectral matched filter and a simple peak-finding algorithm, we have used the 95 and 150 GHz survey data to identify 677 cluster candidates above a signal-to-noise threshold of $\xi = 4.5$. From simulated data, we have estimated the purity of this sample to be 75%; above a threshold of $\xi = 5$, simulations have indicated that the sample should be 95% pure. In our optical/NIR follow-up data, we identified clear overdensities of similarly colored galaxies in the direction of 516 (76%) of the $\xi \geq 4.5$ cluster candidates and 387 (95%) of the $\xi \geq 5$ cluster candidates, confirming the predictions from simulations. Of these confirmed

⁶¹ <http://www.stsci.edu/hst/campaigns/frontier-fields/>

clusters, 415 were first identified in SPT data, including 250 new discoveries reported in this work.

We have also used the optical/NIR data to estimate photometric redshifts for all of our candidates with clear counterparts, and we have estimated lower redshift limits for the candidates without counterparts. We have combined these measurements with spectroscopic redshifts for 141 clusters in the sample to estimate the redshift distribution of the sample. The median redshift is $z_{\text{med}} = 0.55$, 83 (16%) of the confirmed clusters lie at $z \geq 0.8$, and 37 (7%) lie at $z \geq 1$. Using the framework developed for R13, we report masses using the best-fit ξ -mass relation for a fixed flat Λ CDM cosmology with $\Omega_m = 0.3$, $h = 0.7$ and $\sigma_8 = 0.8$. The typical mass of clusters in the sample is $M_{500c} \sim 3.5 \times 10^{14} M_{\odot} h_{70}^{-1}$, nearly independent of redshift. Work is ongoing to improve the mass calibration of SPT clusters using X-ray, galaxy velocity dispersion, and optical weak lensing measurements. Selected data reported in this work, as well as future cluster masses estimated using these data sets, will be hosted at <http://pole.uchicago.edu/public/data/sptz-clusters>.

SZ-selected samples of galaxy clusters from data with sufficient angular resolution are expected to have a nearly redshift-independent mass limit, and the distribution in mass and redshift of the sample presented here is fully consistent with this expectation. This combination of clean selection, large redshift extent, and high typical mass make this sample of particular interest for cosmological and cluster physics analyses.

The catalog presented in this work represents the complete sample of clusters detected at high significance in the 2500 deg² SPT-SZ survey. The program of galaxy cluster science with the SPT continues with the currently fielded SPTpol receiver (Austermann et al. 2012) and will expand further with the expected deployment of the SPT-3G receiver (Benson et al. 2014).

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088.A-0796(A), 088.A-0889(A,B,C), and 286.A-5021 and Gemini programs GS-2009B-Q-16, GS-2011A-C-3, GS-2011B-C-6, GS-2012A-Q-4, GS-2012A-Q-37, GS-2012B-Q-29, GS-2012B-Q-59, GS-2013A-Q-5, GS-2013A-Q-45, GS-2013B-Q-25 and GS-2013B-Q-72 were included in this work. Additional data were obtained with the 6.5 m Magellan Telescopes and the Swope Telescope, which are located at the Las Campanas Observatory in Chile and the MPG/ESO 2.2 m and ESO NTT located at La Silla Facility in Chile. This work is based in part on observations made with the *Spitzer Space Telescope* (PIDs 60099, 70053, 80012 and 10101), which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA. Support for this work was provided by NASA through an award issued by JPL/Caltech. This work is also partly based on observations made with the NASA/ESA *Hubble Space Telescope*, obtained at the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS 5-26555; these observations are associated with programs 12246, 12477, and 13412. The Digitized Sky Surveys were produced at the Space Telescope Science Institute under U.S. Government grant NAG W-2166. The images of these surveys are based on photographic data obtained using the Oschin Schmidt Telescope on Palomar Mountain and the UK Schmidt Telescope. The plates were processed into the present compressed digital form with the permission of these institutions.

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