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Forecasts on interacting dark energy with standard sirens

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We present the predictions with standard sirens at gravitational-wave detectors, such as the Laser Interferometer Space Antenna (LISA) and the Einstein Telescope (ET), for interacting dark energy theories. We focus on four models characterized by couplings between the dark energy field and the dark matter fluid arising from conformal or disformal transformations of the metric, along with an exponential self-interacting potential. To this purpose we construct mock catalogs and perform a Markov chain Monte Carlo analysis by considering ET and LISA standard sirens, and also their combination with baryon acoustic oscillation (BAO) and supernovae Ia (SNIa) data. We find that in all four models considered, the accuracy on the H_0 parameter increases by 1 order of magnitude at 1σ when compared to the SNIa + BAO data set, possibly shedding light in the future on the origin of the H_0 tension. The combination of standard sirens with SNIa + BAO allows to improve the accuracy on some coupling and exponential parameters, hinting at future prospects for constraining interactions in the dark sector.

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I. INTRODUCTION

Understanding the nature of the dark sector of the Universe is one of the greatest endeavors of cosmology at present. This comprises the weakly interacting dark matter (DM)—responsible for the formation and dynamics of structures in the Universe—and dark energy (DE)—the driver of the late-time cosmic acceleration. Together these components dominate about 95% of the energy budget of the Universe. In the Λ -cold-dark-matter (Λ CDM) scenario, i.e. the standard model of cosmology, DE is portrayed simply as a cosmological constant, Λ . This model comes with some theoretical issues [1–3], among which is the fact that fundamental theories do not properly account for the currently measured small value of the cosmological constant. Λ CDM also requires a primordial inflationary period to explain the geometrical flatness, cosmic microwave background (CMB) smoothness, and initial conditions for large-scale structures. More recently, observational tensions on the value of the cosmological parameters H_0 [4–10] and σ_8 [11–16], measured by early- and late-Universe probes, increased the motivation to investigate alternative models of gravity [17].

Consequently, alternative theories are explored by cosmologists in which Λ is promoted to a dynamical DE scalar field, ϕ , namely *quintessence* [18,19] (see [20] for a review), which evolves in time according to its self-interaction potential. While in the standard Λ CDM scenario the two dark components do not directly couple with each other, in a dynamical DE model one can instead consider that they experience some *nonminimal* interaction. Such constructed models are referred to as *coupled quintessence models* [21,22]. The dynamics of the field, along with the dark interaction, could provide a more natural explanation of the accelerated expansion, while also addressing the observational tensions [23]. Nevertheless, the coupling can be formalized at the Lagrangian level through what is known as a conformal/disformal transformation of the metric tensor [24–30]. If this transformation depends directly on the quintessence field, then this is physically equivalent to considering that the DM particles propagate on the geodesics of the transformed metric, $\bar{g}_{\mu\nu}$. In the conformal case, this is achieved from a rescaling of the metric and, consequently, of time- and space-like norms and intervals alike, while preserving the light cones:

$$\bar{g}_{\mu\nu} = C(\phi)g_{\mu\nu}, \quad (1)$$

where C is the conformal function. These find important applications in modified gravity theories as they preserve the structure of scalar-tensor theories of the Brans-Dicke form [31]. Alternatively, one can consider that the metric

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transformation should depend on the first-order partial derivatives of the scalar field as well. This results in a disformal transformation:

$$\bar{g}_{\mu\nu} = C(\phi, X)g_{\mu\nu} + D(\phi)\partial_\mu\phi\partial_\nu\phi, \quad (2)$$

where C and D are the conformal and disformal functions, respectively and $X = -g^{\mu\nu}\partial_\mu\phi\partial_\nu\phi/2$. This gives rise to a more intricate scenario, with a distortion of the metric defined directionally according to the gradient of ϕ . First introduced by Bekenstein [25], the disformal transformations resurged in the cosmological literature [27] when, in analogy to the conformal case, it was shown that they preserve a more general class of scalar-tensor theories categorized in the Horndeski Lagrangian [28]. Disformal transformations in cosmology arise naturally in brane-world models [32,33] and have been the focus of many theoretical proposals for the nature of the dark sector and their interactions [26,34–42]. The stability conditions for the functions C and D have been discussed in Refs. [26–28,40] and the case in which $D \equiv D(\phi, X)$ has been discussed in Refs. [43,44].

The coupling in the dark sector gives rise to an additional gravitational fifth force in the Universe between DM particles, mediated by the DE field. This new force leaves distinct features in the background equations, as well as signatures in the cosmological density perturbations that describe the formation of structures [45]. Although these deviations from the benchmark model are constrained to be small (especially at the background level), they are still expected to leave detectable, characteristic observational imprints that the data can probe. These are essential to test the viability of such alternative models by identifying the range of validity of the parameter space and the robustness of its predictions.

In the past few years, we have witnessed the rise of gravitational-wave (GW) astronomy as a new independent probe of gravitational effects [46]. An accurate redshift-luminosity relation can be constructed when GW events are combined with an electromagnetic (EM) counterpart multimessenger signal. These observations become *standard sirens* [47], analogous to the standard candles used in local EM measurements. So far, only one GW event, GW170817, with a corresponding EM counterpart, GRB170817A, has been detected by the Laser Interferometer Gravitational-Wave Observatory (LIGO)-Virgo and International Gamma-ray Astrophysics Laboratory–Fermi collaborations, respectively, and which originated from the merger of a binary pair of neutron stars [48,49]. This single combined detection had a strong impact on the allowed modifications to the gravitational interaction by ruling out many proposals [50–54] with many other models further constrained [55–59].

Current GW detectors, (Advanced) Virgo [60], (Advanced) LIGO [61] and the Kamioka Gravitational

Wave Detector [62], are second-generation (2G) ground-based detectors, with another one under planning (2030), the Indian Initiative in Gravitational-wave Observations [63]. The increasing number of detectors will boost the capabilities of GW astronomy both in the number of confirmed events (a larger volume of the Universe is covered) and sky localization (a better triangulation of the source), which will also aid in the search for a counterpart. However, 2G detectors are limited in their sensitivity and future third-generation (3G) ground-based detectors are designed to become more sensitive, precise and capable of probing a larger range of frequencies. Special emphasis should be given to the Einstein Telescope (ET), which is expected to improve the current sensitivity by a factor of 10 [64]. ET will also extend the redshift range, e.g. $z \sim 5$ for binary black holes compared to $z \sim 0.5$ for 2G detectors [65]. The number of detectable multimessenger events is expected to reach tens of thousands of standard sirens [66]. While these ground-based detectors will cover a frequency band in the range $1 \lesssim f \lesssim 10^3$ Hz [67], the upcoming space-based 3G detectors, such as the Laser Interferometer Space Antenna (LISA) [68] will have a peak sensitivity near 10^{-3} Hz and will be able to detect GW events beyond $z = 20$, probing a wide range of targets. There are many proposals of 3G GW observatories, such as the DECi-hertz Interferometer Gravitational wave Observatory [69]. However, we have opted to focus our analysis on ET and LISA covering ground- and space-based experiments.

In this paper, we aim at forecasting the constraining power of future 3G detectors; in particular we focus on ET and LISA. Given the potential of such missions we are interested in assessing their ability to constrain modifications to general relativity as well as to provide complementary constraints on H_0 using standard sirens. We investigate four models characterized by coupling functions between the DM and DE fields: a conformally coupled quintessence field [21,22], characterized by a conformal coupling in the form of an exponential function of the scalar field; a kinetic model [70] with a conformal function given by a power law of the kinetic term of the scalar field; a purely disformally coupled quintessence field [29,38] with a constant disformal coupling; and a mixed disformally coupled quintessence [42,43] which combines the previous model with an exponential conformal coupling. All the scenarios considered are characterized by the same simple exponential potential which introduces one more free parameter. The models we consider differ considerably in the way the (effective) coupling between DM and DE evolves and, in particular, they differ in their background evolution. As such, they represent a well-studied sample of models of interacting dark energy suitable for our analysis. We construct our pipeline following the methodology presented in [71–74] providing for the first time GW

forecasts on the free parameters of the four models in question.

This paper is organized as follows. We start by giving a brief introduction to the physics of standard sirens in Sec. II. Section III provides an overview of the methodology used and the details on the simulation of the standard siren events developed for this study, as well as a brief account of the data set combinations considered. We outline the criteria for particular catalog choices, and discuss the sampling method employed for the forecasts. In Sec. IV we introduce each of the four models under study and present the results of our analysis, emphasizing their significant implications. Last, in Sec. V, we summarize our results and outline our concluding thoughts and future prospects.

II. GRAVITATIONAL WAVES AS STANDARD SIRENS

Interferometers are sensitive to the strain, $h(t)$ from a GW event, which in the transverse-traceless gauge is described as, [71]

$$h(t) = F_{\times}(\theta_0, \phi_0, \psi)h_{\times}(t) + F_{+}(\theta_0, \phi_0, \psi)h_{+}(t), \quad (3)$$

where θ_0 and ϕ_0 define the initial location of the event relative to the detector in polar coordinates, ψ is the polarization of the GW event, and t is cosmic time. We adopt a random sampling method in the range $[0 - 2\pi]$ for θ_0 and $[0 - \pi]$ for both ϕ_0 and ψ . The factors $F_{\times,+}$ describe the antenna beam pattern function,

$$\begin{aligned} F_{\times}^{(1)} &= \frac{\sqrt{3}}{2} \left[\frac{1}{2} (1 + \cos^2(\theta)) \cos(2\phi) \cos(2\psi) + \cos(\theta) \sin(2\phi) \cos(2\psi) \right], \\ F_{+}^{(1)} &= \frac{\sqrt{3}}{2} \left[\frac{1}{2} (1 + \cos^2(\theta)) \cos(2\phi) \cos(2\psi) - \cos(\theta) \sin(2\phi) \cos(2\psi) \right]. \end{aligned} \quad (4)$$

The superscript number indicates which interferometer is being considered, e.g. LISA only has two separate interferometers and therefore $F^{(3)} = 0$. Since the detectors are spatially distributed in an equilateral triangle formation, the other two antenna pattern functions relate to $F_{\times,+}^{(1)}$ as

$$F_{\times,+}^{(1)}(\theta, \phi, \psi) = F_{\times,+}^{(2)}\left(\theta, \phi + \frac{2\pi}{3}, \psi\right) = F_{\times,+}^{(3)}\left(\theta, \phi + \frac{4\pi}{3}, \psi\right). \quad (5)$$

As LISA is sensitive to lower frequencies, and equivalently larger masses, it can detect GW events of inspiral mergers lasting over several months, during which the interferometer's position will change relative to the event. This change in position is accounted for following the method described in [72]. The time scale of the event is described as

$$t = t_c - 5(8\pi f)^{-8/3} M_c^{-5/3}. \quad (6)$$

Here t_c is the time of the merger, t indicates the time at which LISA detects the merger, f is the frequency of the GW, and M_c is the chirp mass. The location angles are updated accordingly:

$$\theta = \cos^{-1} \left[\frac{1}{2} \cos(\theta_0) - \frac{\sqrt{3}}{2} \sin(\theta_0) \cos\left(\frac{2\pi t}{T} - \phi_0\right) \right], \quad (7)$$

$$\phi = \frac{2\pi t}{T} - \tan \left[\frac{\sqrt{3} \cos(\theta_0) + \sin(\theta_0) \cos\left(\frac{2\pi t}{T} - \phi_0\right)}{2 \sin(\theta_0) \cos\left(\frac{2\pi t}{T} - \phi_0\right)} \right], \quad (8)$$

which, in turn, are used to update the beam pattern functions. Here we have specified the period, T , as the orbit around the Sun.

While the individual masses of the objects are not directly discernible, GW detectors are sensitive to the chirp mass, a collective mass quantity related to the frequency evolution of the signal emitted before the merger, during the inspiral phase of the binary [75], defined as

$$M_c = (1+z) \left(\frac{(m_1 m_2)^3}{m_1 + m_2} \right)^{1/5}, \quad (9)$$

where $(1+z)$ is a conversion redshift factor from the physical to the observational chirp mass.

The Fourier transform of the strain using the stationary phase approximation [73] reads

$$\mathcal{H} = \mathcal{A}f^{-7/6}e^{i\Psi(f)}, \quad (10)$$

where $\Psi(f)$ is the phase of the waveform. Notice that when \mathcal{H} is inserted into Eq. (14), the exponential term disappears, meaning that the $\Psi(f)$ factor can be discarded for this analysis. \mathcal{A} is the Fourier amplitude of the waveform,

$$\mathcal{A} = \frac{M_c^{5/6}}{d_L} \pi^{-2/3} \sqrt{\frac{5}{96}} \sqrt{[F_+(1 + \cos^2(l))]^2 + (2F_\times \cos(l))^2}, \quad (11)$$

where d_L is the luminosity distance from the merger and l is the inclination angle, which we sample randomly between $[0^\circ, 20^\circ]$, as that is the maximum detection inclination range.

LISA has been designed to effectively measure frequencies as low as $f_{\min} = 1 \times 10^{-4}$ Hz, which is why it stands as a promising probe of extreme mass ratio inspiral (EMRI) and binary massive black hole (BMBH) mergers. For the purpose of the simulations, the upper bound frequency of LISA is determined by two quantities: the structure of LISA itself and the last stable orbit of the merging system. LISA can detect frequencies up to $f_{\max} = c(2\pi L)^{-1}$, where L is the length of LISA’s interferometer arm, taken to be 2.5 Gm and c is the speed of light. Moreover, the total mass of an orbiting system is inversely proportional to its measured frequency, implying that even though massive mergers give rise to large detection amplitudes, the frequency will fall below f_{\min} . Therefore if the last stable orbit frequency, $f_{\text{LSO}} = (6^{3/2}2\pi M_{\text{obs}})^{-1}$, with M_{obs} being the observed total mass, is found to be lower than f_{\min} , we disregard that simulated event. If otherwise it lies between f_{\min} and f_{\max} then f_{LSO} becomes the new maximum frequency for that event.

III. METHODOLOGY AND DATA SETS

Given the main objective of this study, we create simulated data that forecasts the potential future observations of standard siren events. Specifically, we focus on those that could be detected by ET and LISA. Below, we provide a concise overview of the samples we have generated along with the methodology and the data combinations used.

A. Simulated cosmology

To simulate GW catalogs from future probes of black hole mergers, the following cosmological quantities are required: the redshift of the merger, z , the value of the Hubble rate at merger, $H(z)$, its comoving and luminosity distance, $d_c(z)$ and $d_L(z)$ respectively, and the cosmic time between the merger and measurement, t . For this purpose, we resort to the public Einstein-Boltzmann code CLASS¹ [76–78], which we extend to accommodate general models of interacting dark energy. This new patch is then used to

provide a *mock Universe* adopting a flat Λ CDM as the fiducial model to simulate the GW data, according to the best-fit cosmological parameters of the *Planck* 2018 data release [5]. These are: the Hubble parameter at present time, $H_0 = 67.32 \text{ km s}^{-1} \text{ Mpc}^{-1}$, the density of baryons, $\Omega_b h^2 = 0.022383$ (with $h = H_0/100$) and the density of cold dark matter, $\Omega_c h^2 = 0.12011$. Furthermore, we are also interested in the derived quantity $\Omega_m^0 = \Omega_b + \Omega_c$, which for the fiducial *Planck* case is $\Omega_m^0 = 0.3144$.

Provided with the background cosmology, we simulate the merger events to determine the redshift-luminosity relation. First, we generate a redshift distribution of events weighted by a probability distribution. The characteristics of these events, such as the chirp mass, are simulated using a uniform distribution. Although each instance of running the script will yield a different set of simulated data, the resulting conclusions will be unaltered as the fiducial parameters constrain the mock data. Once the mergers have been simulated, we emulate the measurement process from the inspiral, yielding the errors associated with each event. As such, simulated data points are removed if they produce a signal-to-noise ratio below the threshold.

B. Distribution of simulated merger events

ET is designed to probe a range of frequencies, f , similar to that of LIGO, thereby probing merger events of nearby compact objects such as binary neutron stars (BNSs) in the mass range of $[1, 2]$, $[1, 2]M_\odot$, and black hole–neutron star binaries (BHNSs) in the mass range $[3, 10]$, $[1, 2]M_\odot$, with the $[\cdot, \cdot]$ notation indicating the uniformly distributed mass ranges considered. Advanced LIGO claims a ratio of BHNS to BNS merger events of ~ 0.03 [79]. The redshift probability distribution of these events is proportional to

$$P \propto \frac{4\pi d_c(z)R(z)}{(1+z)H(z)}, \quad (12)$$

where the comoving distance and the Hubble parameter are taken at various redshifts determined by CLASS. $R(z)$ stands for the merger rate, which, at a linear approximation level, is [74]

¹https://github.com/lesgourg/class_public.

$$R = \begin{cases} 1 + 2z & \text{if } z < 1, \\ \frac{3}{4}(5 - z) & \text{if } 1 \leq z < 5, \\ 0 & \text{otherwise.} \end{cases} \quad (13)$$

On the other hand, LISA will target lower frequencies when compared with other proposed 3G detectors, implying sensitivity to events from larger mass binary systems since $f \propto M^{-1}$. Therefore we focus on simulating the detection of events from EMRIs and BMBHs in the ranges $[1 - 30]$, $[10^4 - 10^8]M_\odot$ [80] and $[10^4 - 10^8]$, $[10^4 - 10^8]M_\odot$ [81], respectively. The number of detected BMBH to EMRI events is estimated to follow a 2:1 ratio according to the mission's proposal [82,83].

Although in principle LISA will also be able to probe mergers of binary intermediate-mass black holes (IMBHs) and binary compact objects, we opt to discard these from the simulations. This is due to the fact that there is no definitive observational proof of IMBHs, and expected events from binary compact objects will only be observed at redshifts $z \approx 3$ [84]. These events are insignificant since we are interested in the higher range of redshifts for our cosmology.

Considering events involving BMBHs only, the redshift probability distributions are based on the histogram for the L6A2M5N2 mission specification [81] which considers three formation processes of BMBHs. We consider the light seed model (pop III) which attributes the formation of BMBHs to the remnants of population III stars around $z = 15-20$. In [81], two additional scenarios for massive black hole formation were investigated, namely delay and no delay scenarios. These cases involve the collapse of gas in a galactic center at $z = 15-20$, leading to the formation of a black hole through a heavy seed mechanism with and without a delay between galaxy merger and the merger of the central massive black hole. Further information on these scenarios can be found in Ref. [85].

In our investigation we provide mock data and obtain forecasts for both the delay and no delay cases. However, the analysis reveals that the predicted constraining power from these models shows no actual improvement compared to the pop III case. Consequently, in this paper, we focus solely on the pop III model, as it proves sufficient to forecast the constraining power of LISA.

C. Simulation of measurements and errors

To simulate the errors associated with the standard siren catalog, we follow the methodology of [71–74]. An apparent detection of a GW event is assessed by evaluating the signal-to-noise ratio (SNR), ρ , and only confirmed if $\rho > 8$. The SNR is defined as

$$\rho_{1,2,3}^2 = 4 \int_{f_{\min}}^{f_{\max}} df \frac{|\mathcal{H}|^2}{S_h}, \quad (14)$$

where the number labels indicate the interferometer being considered. \mathcal{H} has been defined in Eq. (10) and S_h is the noise power spectral density, an SNR weighting function that accounts for the particular properties of the instruments used. For ET, in particular, S_h is designed to follow

$$S_h^{(\text{ET})} = S_0 \left(x^{p_1} + a_1 x^{p_2} + a_2 \frac{1 + \sum_{n=1}^6 b_n x^n}{1 + \sum_{m=1}^4 c_m x^m} \right), \quad (15)$$

where, $x = f/200 \text{ Hz}^{-1}$, $S_0 = 1.449 \times 10^{-52} \text{ Hz}$, $p_1 = -4.05$, $p_2 = -0.69$, $a_1 = 185.62$, $a_2 = 232.56$, $b_n = \{31.18, -64.72, 52.24, -42.16, 10.17, 11.53\}$, and $c_m = \{13.58, -36.46, 18.56, 27.43\}$, assuming a lower cutoff at $f = 1 \text{ Hz}$. On the other hand, for LISA, S_h depends on the instrumental (or short) noise, S_{inst} , the noise from low-level acceleration, S_{acc} , and the confusion background noise, S_{conf} [85]:

$$S_h^{(\text{LISA})} = \frac{20}{3} \frac{4S_{\text{acc}} + S_{\text{inst}} + S_{\text{conf}}}{L^2} \left[1 + \left(\frac{fL}{0.81c} \right) \right], \quad (16)$$

where $S_{\text{acc}} = 9 \times 10^{-30} / (2\pi f)^4 (1 + 10^{-4}/f)$, $S_{\text{inst}} = 2.22 \times 10^{-23}$ and $S_{\text{conf}} = 2.65 \times 10^{-23}$. Therefore, the total SNR contribution for each detector is given by combining (10) with either Eq. (15) or (16) for the ET and LISA, respectively:

$$\rho_{\text{tot}} = \sqrt{\rho_1^2 + \rho_2^2 + \rho_3^2}. \quad (17)$$

The instrumental error in the luminosity distance is determined *via* the Fisher matrix,

$$\sigma_{d_L}^{\text{inst}} \approx \left\langle \frac{\partial \mathcal{H}}{\partial d_L}, \frac{\partial \mathcal{H}}{\partial d_L} \right\rangle^{-\frac{1}{2}}, \quad (18)$$

following [72]. Since $\mathcal{H} \propto d_L^{-1}$ this results simply in

$$\sigma_{d_L}^{\text{inst}} \approx \frac{2d_L}{\rho}, \quad (19)$$

where the factor of 2 accounts for the symmetry in the inclination angle, which actually ranges from -20° to 20° . The error due to gravitational lensing is,

$$\sigma_{d_L}^{\text{len}} = \frac{d_L}{2} \times 0.066 [4(1 - (1+z)^{1/4})]^{1.8}, \quad (20)$$

reduced by a half to account for both the merger and ringdown of the event.

Being space-based, LISA is also subject to an error associated with the peculiar velocities of GW sources [86]:

$$\sigma_{d_L}^{\text{pec}} = d_L \frac{\sqrt{\langle v^2 \rangle}}{c} \left[1 + \frac{c(1+z)}{Hd_L} \right], \quad (21)$$

with an estimate of the peculiar velocity of the host galaxy with respect to the Hubble flow of $\sqrt{\langle v^2 \rangle} = 500 \text{ km s}^{-1}$.

Bringing all the contributions together, the total error in the luminosity distance is simply a combination of the errors in Eqs. (19)–(21):

$$\sigma_{d_L} = \sqrt{(\sigma_{d_L}^{\text{inst}})^2 + (\sigma_{d_L}^{\text{len}})^2 + (\sigma_{d_L}^{\text{pec}})^2}. \quad (22)$$

The simulation allows us to interpolate any number of events over a continuous redshift distribution in the range $0 < z \lesssim 5$ for ET and $0 < z \lesssim 10$ for LISA. However, the number of mergers detected by ET will depend on factors such as running costs and the complementary detection with other experiments [71]. ET is expected to report more than 10^4 mergers yearly. However, due to the scarcity of EM counterpart signals, the predicted number of detectable mergers with an actual EM counterpart over the course of 10 years is approximately 200 [87]. According to [81], LISA’s number of detected mergers, for a 10-year mission proposal, is 56 events.

To incorporate uncertainty into the luminosity distance of each merger, we apply a Gaussian distribution centered around the background cosmology. The standard deviation for this distribution is set to the calculated errors, σ_{d_L} . This introduces artificial randomness around each merger, leading to a larger deviation from Λ CDM in LISA compared to ET. The reason for this difference is that LISA probes larger redshifts, which are associated with larger errors, resulting in a broader spread of the data, as depicted in Fig. 1.

D. Data sets and likelihoods

To examine the fit of the simulated data to the coupled quintessence models considered in this study, we employ the Markov chain Monte Carlo (MCMC) method, using samples generated from our modified version of CLASS interfaced

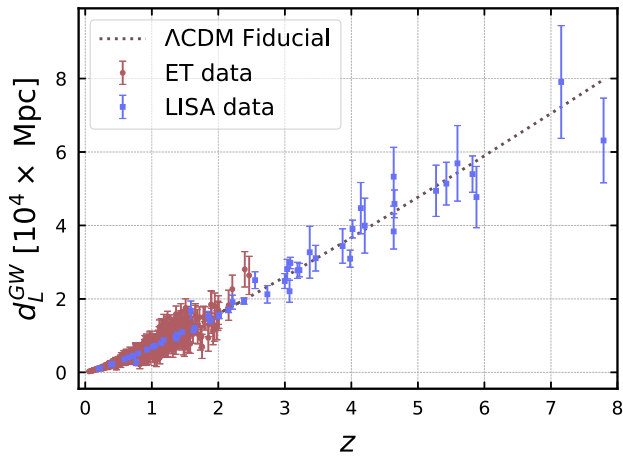


FIG. 1. Mock data from ET (red circle markers) and LISA (blue squared markers), for the fiducial model, Λ CDM, shown by the grey dotted line.

with the Monte Python Sampler [88,89]. In particular, we resort to the nested sampling algorithm through the MultiNest² [90–92] and PyMultiNest³ [93] packages to estimate observational constraints on the free parameters, instead of the traditional Metropolis-Hastings algorithm. The Metropolis-Hastings algorithm struggles to explore the full line of degeneracies between the parameters, resulting in false peaks in the posterior distribution, which the sampler cannot move away from. Nested sampling is able to explore the full extent of the degeneracies as it is much better suited for multimodel sampling (see Sec. IV) and other more complicated distributions. Subsequently, we analyze the MCMC chains and present the results using the GetDist⁴ Python package [94].

The likelihood function for the simulated data set of standard siren GW events is constructed according to the effective Gaussian distribution:

$$\ln \mathcal{L}_{SS} = -\frac{1}{2} \sum_{i=1}^n \left[\frac{d_{SS}^{(\text{obs})}(z_i) - d_{SS}(z_i)}{\sigma_{d_{L,i}}} \right]^2, \quad (23)$$

where $d_{SS}^{(\text{obs})}(z)$ is the observed luminosity distance, which in this case corresponds to the samples generated according to the procedure outlined above; $d_{SS}(z)$ is the model-dependent theoretical prediction for the luminosity distance of the event, computed numerically with the modified CLASS code; σ_{d_L} is the total error in the luminosity distance, as defined in Eq. (22); and n is the number of observed events.

Since we want to forecast the constraining power of standard siren data probed by ET and LISA on coupled quintessence models, we assess the independent and combined constraints with *current* background data. This allows for a direct comparison of whether GW catalogs will improve the constraints on $\{\Omega_m^0, H_0\}$ and on the model-specific parameters, affecting the background evolution. In particular, we include baryonic acoustic oscillation (BAO) data from the Sloan Digital Sky Survey (SDSS) DR7 Main Galaxy Sample [95], SDSS DR12 consensus release [96] and the 6dF Galaxy Survey [97], in combination with distance *moduli* measurements of 1048 type Ia supernova (SNIa) data from Pantheon [98]. This combined data set is referred to as “SNIa + BAO.”

Our analysis involves a set of free sampling parameters, including the baseline Λ CDM cosmological parameters (Ω_m^0, H_0) and the parameters associated with each coupled quintessence model. For the latter, we take the corresponding Λ CDM limit case as fiducial value. The models discussed in Sec. IV reduce to Λ CDM in the following limits: $\lambda = 0$ and $\beta = 0$ for Sec. IV A; $\lambda = 0$ and $\alpha = 0$ for

²<https://github.com/farhanferoz/MultiNest>.

³<https://github.com/JohannesBuchner/PyMultiNest>.

⁴<https://github.com/cmbant/getdist>.

TABLE I. Flat priors on the cosmological and model parameters sampled in Sec. IV.

Model	Parameter	Prior
All	$\Omega_b h^2$	[0.018, 0.03]
	$\Omega_c h^2$	[0.1, 0.2]
	h	[0.6, 0.8]
	λ	[0, 2]
Sections IV A and IV D	β	[0, 2]
Section IV B	α	[0, 0.001]
Sections IV C and IV D	D_0/meV^{-1}	[0, 2]

Sec. IV B; $\lambda = 0$ and $D_0 = 0$ for Sec. IV C; $\lambda = 0$, $\beta = 0$ and $D_0 = 0$ for Sec. IV D. We adopt flat priors for all parameters within the ranges specified in Table I.

IV. FORECAST RESULTS

In what follows, we employ the methodology and the data sets discussed in Sec. III to investigate the power that LISA and ET standard sirens have in constraining the cosmological parameters $\{\Omega_m^0, H_0\}$, the model-dependent conformal and disformal coupling parameters and the steepness of the self-interacting potential. In particular, we consider four interacting DE models: a standard coupled quintessence model, a kinetically coupled model, a constant disformal model and a mixed conformal-disformal model. For each of the four scenarios, we provide a brief review of the theoretical framework before presenting the forecasts obtained considering the specifications and assumptions discussed in previous sections.

In each subsection that follows, we show the resulting 2D contours, and 1D marginalized posterior distributions for $\{H_0, \Omega_m^0\}$ plus the set of model-specific parameters (see Table I) for the cases of ET and LISA and their combination. These plots also include a combination of SNIa + BAO, and for reference, the results of SNIa + BAO alone. The results are also summarized in a table with their corresponding 1σ , identified in the text with the notation $\{\sigma_p\}$ where p is an index spanning over the model parameters. We also use $\mathcal{F}_p^{(i,j)} = \{\sigma_p^{(j)}/\sigma_p^{(i)}\}$ where here i and j stand for two different data sets, to denote the change in error for the specific parameter p .

A. Conformal coupling

The first model we consider is the conformal coupling model, for which

$$C(\phi) = e^{2\beta\phi/M_{\text{Pl}}} \quad \text{and} \quad V(\phi) = V_0 e^{-\lambda\phi/M_{\text{Pl}}}, \quad (24)$$

where $C(\phi)$ is defined according to Eq. (1) and $V(\phi)$ is the DE potential energy. The exponential parameters β and λ

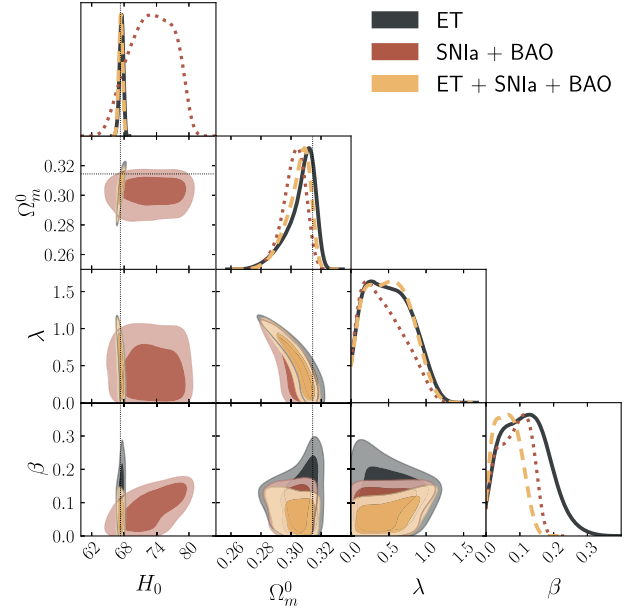


FIG. 2. 68% and 95% confidence limits (C.L.) 2D contours and 1D marginalized posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, \beta\}$ in the conformal coupled quintessence model with the ET mock data (charcoal filled line), SNIa + BAO data (red dotted line) and their combination (yellow dashed line). The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

are constant dimensionless parameters and V_0 is a constant with dimensions of $(\text{mass})^4$ that sets the energy scale of the potential.⁵ In such models, the mass of the DM particles becomes ϕ dependent and the DE field mediates a long-range force between DM particles so that the effective gravitational coupling is given by $G_{\text{eff}} = G_N(1 + 2\beta^2)$ [21,99,100]. The free parameters we are particularly interested in are the slope of the potential λ and the coupling parameter β . Constraints on this model have been obtained in Ref. [101] using background data only $[H(z), \text{BAO and supernova Union2.1}]$. Using these data, the authors found the following upper limits: $\beta < 0.193$ and $\lambda < 1.27$. In [102] stronger constraints have been obtained using *Planck* data, BAO and SNIa data, also in line with Ref. [103], in which the authors found $\beta < 0.0298$ and $\lambda < 0.6$ for the 1σ upper limits.

According to the results in Figs. 2–4, summarized in Table II, we comment on the resulting constraints for GW data sets compared with SNIa + BAO for the parameters $\{\Omega_m^0, H_0, \beta, \lambda\}$. When ET standard sirens are considered, we find that the cosmological and model parameters can be

⁵For numerical purposes and to avoid degeneracies, V_0 is not taken to be a free parameter. Instead, it serves as a shooting parameter to set the fiducial value of Ω_ϕ^0 fulfilling the flatness condition.

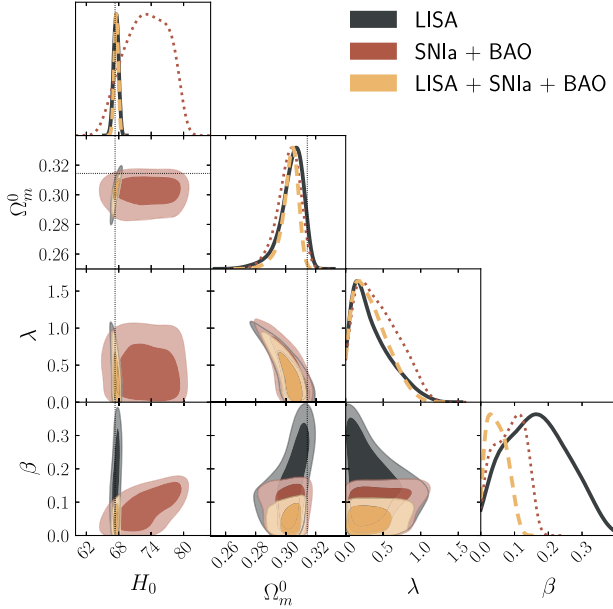


FIG. 3. 68% and 95% C.L. 2D contours and 1D marginalized posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, \beta\}$ in the conformal coupled quintessence model with LISA mock data (charcoal filled line), SNIa + BAO data (red dotted line) and their combination (yellow dashed line). The scale is the same as in Fig. 2 for comparison purposes, with the SNIa + BAO contours standing as the reference. The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

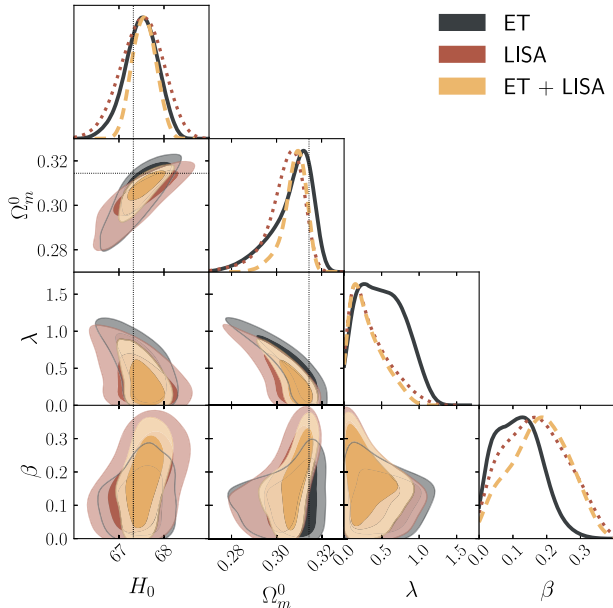


FIG. 4. 68% and 95% C.L. 2D contours and 1D marginalized posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, \beta\}$ in the conformal coupled quintessence model with ET mock data (charcoal filled line), LISA mock data (red dotted line) and their combination (yellow dashed line). The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

constrained at 1σ with an accuracy $\{0.0080, 0.37, 0.0070, 0.32\}$ for ET alone and $\{0.0075, 0.36, 0.039, 0.31\}$ for ET + SNIa + BAO, resulting in a change in error of $\mathcal{F}_{\Omega_m^0, H_0, \beta, \lambda}^{(\text{ET}, \text{ET} + \text{SNIa} + \text{BAO})} = \{0.94, 0.97, 0.56, 0.97\}$. Thus, the forecasted constraints of ET + SNIa + BAO, compared to ET alone, have increased accuracy in all parameters shown by the reduction in σ . This trend is also present in the LISA data set with the cosmological and model parameters constrained with an accuracy of $\{0.0071, 0.47, 0.098, 0.24\}$ for LISA alone and $\{0.0051, 0.37, 0.031, 0.22\}$ for LISA + SNIa + BAO, resulting in a reduction in σ by a factor of $\mathcal{F}_{\Omega_m^0, H_0, \beta, \lambda}^{(\text{LISA}, \text{LISA} + \text{SNIa} + \text{BAO})} = \{0.72, 0.79, 0.32, 0.92\}$. For the combination of just SNIa + BAO we obtain an accuracy of $\{0.0074, 4.1, 0.049, 0.28\}$. Compared to ET + SNIa + BAO and LISA + SNIa + BAO to SNIa + BAO, there is also a reduction in σ for all parameters except one. The ET + SNIa + BAO data set results in a nominal increase in $\sigma_{\Omega_m^0}$ compared to SNIa + BAO. Comparing the errors of ET and LISA to SNIa + BAO we see only minor changes in the constraining power regarding Ω_m^0 , with ET performing slightly worse and LISA slightly better. A similar trend occurs for the parameter λ , with ET performing nominally worse and LISA better. However, particular attention should be given to the significant reduction in σ_{H_0} when comparing ET and LISA to SNIa + BAO. There is a reduction in the error by a factor of $\mathcal{F}_{H_0}^{(\text{SNIa} + \text{BAO}, \text{ET})} = 0.090$ and $\mathcal{F}_{H_0}^{(\text{SNIa} + \text{BAO}, \text{LISA})} = 0.11$. Forecasting GWs will improve the constraints on H_0 , suggesting that GWs will be critical in addressing the Hubble tension. On the other hand, we see the opposite effect with β , with an increase in the error by a factor of $F_{\beta}^{(\text{SNIa} + \text{BAO}, \text{ET})} = 1.4$ and $\mathcal{F}_{\beta}^{(\text{SNIa} + \text{BAO}, \text{LISA})} = 2.0$. Nonetheless, when the background data is combined with ET and/or LISA, the constraints improve by $F_{\beta}^{(\text{ET}, \text{ET} + \text{SNIa} + \text{BAO})} = 0.56$ and $\mathcal{F}_{\beta}^{(\text{LISA}, \text{LISA} + \text{SNIa} + \text{BAO})} = 0.32$. In comparing the constraining power of ET and LISA (see Fig. 4), it is evident that they have comparable spreads for the cosmological parameters. An interesting feature we observe is that ET is more constraining in regard to H_0 . We attribute this feature to the fact that the ET catalog has more data points than LISA at low redshifts, as illustrated in Fig. 1.

By combining GW data from LISA and ET, which implies an increase of data points over a wide range of redshifts, we predict an enhanced constraining power in the cosmological parameters, $\{H_0, \Omega_m^0\}$, compared to the SNIa + BAO case, more precisely $\mathcal{F}_{\Omega_m^0, H_0}^{(\text{SNIa} + \text{BAO}, \text{ET} + \text{LISA})} = \{0.65, 0.063\}$. However, for the model parameters β and λ , we observe modifications to the constraining power with $\mathcal{F}_{\beta, \lambda}^{(\text{SNIa} + \text{BAO}, \text{ET} + \text{LISA})} = \{1.8, 0.86\}$. The combination of ET + LISA with SNIa + BAO results in a negligible

TABLE II. Marginalized constraints on cosmological and model parameters for the conformal coupled quintessence model at 68% C.L.

Conformal coupled quintessence								
Data sets	Ω_m^0	$\sigma_{\Omega_m^0}$	H_0	σ_{H_0}	β	σ_β	λ	σ_λ
SNIa + BAO	$0.3019^{+0.0088}_{-0.0059}$	0.0074	$73.2^{+4.7}_{-3.5}$	4.1	$0.085^{+0.055}_{-0.043}$	0.049	$0.42^{+0.20}_{-0.36}$	0.28
ET	$0.307^{+0.011}_{-0.0050}$	0.0080	$67.49^{+0.39}_{-0.34}$	0.37	$0.115^{+0.060}_{-0.079}$	0.070	$0.50^{+0.26}_{-0.38}$	0.32
ET + SNIa + BAO	$0.3046^{+0.0099}_{-0.0051}$	0.0075	67.37 ± 0.36	0.36	$0.063^{+0.033}_{-0.045}$	0.039	$0.49^{+0.26}_{-0.35}$	0.31
LISA	$0.3039^{+0.0093}_{-0.0049}$	0.0071	$67.50^{+0.50}_{-0.44}$	0.47	$0.167^{+0.085}_{-0.11}$	0.098	$0.33^{+0.15}_{-0.32}$	0.24
LISA + SNIa + BAO	$0.3028^{+0.0065}_{-0.0036}$	0.0051	67.52 ± 0.37	0.37	$0.048^{+0.025}_{-0.037}$	0.031	$0.33^{+0.15}_{-0.29}$	0.22
ET + LISA	$0.3079^{+0.0061}_{-0.0034}$	0.0048	67.56 ± 0.26	0.26	$0.178^{+0.099}_{-0.081}$	0.090	$0.30^{+0.13}_{-0.27}$	0.24
ET + LISA + SNIa + BAO	$0.3044^{+0.0063}_{-0.0032}$	0.0048	67.45 ± 0.28	0.28	$0.052^{+0.028}_{-0.038}$	0.033	$0.35^{+0.17}_{-0.30}$	0.24

change of constraining power for Ω_m^0, H_0, λ . Only the parameter β is more constrained when the data sets are combined, with the 1σ reduced by almost a third.

Compared to the current background constraints mentioned in the beginning of the section, we find in our analysis that the upper bounds at 1σ on the model parameters are improved in the following cases: $\beta < 0.14$ and $\lambda < 0.62$ (SNIa + BAO); $\beta < 0.175$ and $\lambda < 0.76$ (ET); $\beta < 0.096$ and $\lambda < 0.75$ (ET + SNIa + BAO); $\lambda < 0.48$ (LISA and LISA + SNIa + BAO) and $\beta < 0.073$ (LISA + SNIa + BAO); $\lambda < 0.43$ (ET + LISA); $\lambda < 0.52$ and $\beta < 0.08$ (ET + LISA + SNIa + BAO).

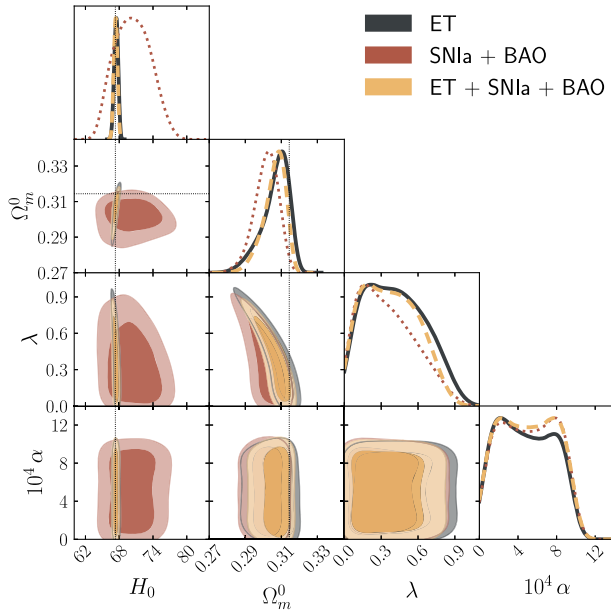


FIG. 5. 68% and 95% C.L. 2D contours and 1D posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, 10^4 \alpha\}$ in the kinetic conformal coupled quintessence model with ET (charcoal filled line), SNIa + BAO (red dotted line) data and their combination (yellow dashed line). The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

B. Kinetic conformal coupling

As an example of a coupled quintessence model in which the conformal function is less trivial, we focus on a pure dependence on derivatives of the scalar field through the kinetic term of ϕ , $X = -\partial_\mu \phi \partial^\mu \phi / 2$, to which we refer as the kinetic coupling. Such a setting has been proposed in [104] (see references therein as well), and we focus on the particular example of a power law, as studied in [70]. Even though this model is proposed based on a Lagrangian framework ($\mathcal{L}_{\text{DM}} \rightarrow (X/M_{\text{Pl}}^4)^\alpha \mathcal{L}_{\text{DM}}$), at the background

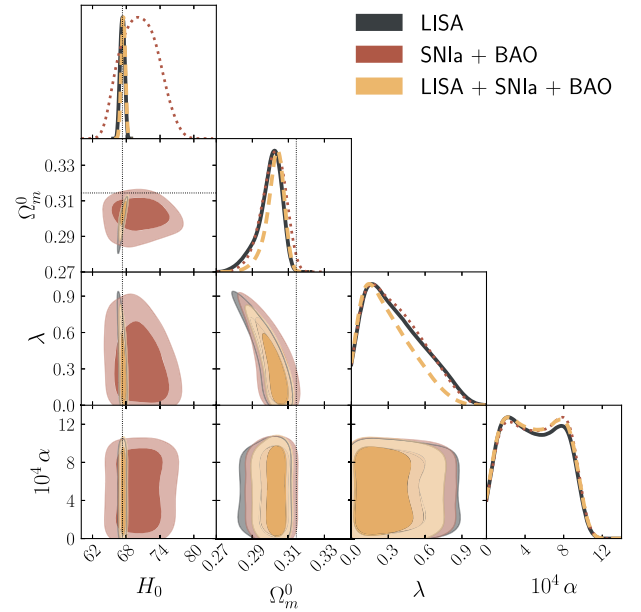


FIG. 6. 68% and 95% C.L. 2D contours and 1D posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, 10^4 \alpha\}$ in the kinetic conformal coupled quintessence model with LISA mock data (charcoal filled line), SNIa + BAO (red dotted line) data and their combination (yellow dashed line). The scale is the same as in Fig. 5 for comparison purposes, with the SNIa + BAO contours standing as the reference. The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

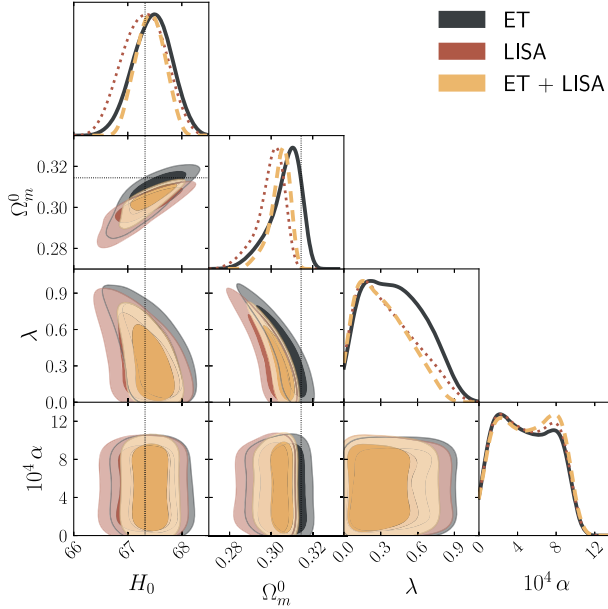


FIG. 7. 68% and 95% C.L. 2D contours and 1D marginalized posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, 10^4\alpha\}$ in the kinetic conformal coupled quintessence model with ET mock data (charcoal filled line), LISA mock data (red dotted line) and their combination (yellow dashed line). The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

level it is equivalent to the kinetic-dependent conformal transformation $\bar{g}_{\mu\nu} = C(X)g_{\mu\nu}$, with

$$C(X) = (M_{\text{pl}}^{-4}X)^{2\alpha} \quad \text{and} \quad V(\phi) = V_0 e^{-\lambda\phi/M_{\text{pl}}}, \quad (25)$$

where α is a dimensionless constant and a simple exponential potential has been assumed just like in the previous case, and the same considerations apply for λ and V_0 .

In summary, an analysis based on *Planck* and the SNIa + BAO background data in Ref. [70] reveals the power of BAO data in constraining Ω_m^0 , which is highly correlated with the steepness of the potential λ . The coupling parameter α is constrained to be of the order of 10^{-4} . The constraints on the cosmological parameters are found

to be compatible with the Λ CDM ones within the errors. Moreover, a positive correlation between H_0 and Ω_m^0 is identified. While this trend is attributed to the evolution of the linear perturbations for nonvanishing α , we find that it is still present for the background standard siren data sets.

From the results presented in Figs. 5–7, and summarized in Table III, we analyze the constraints on the parameters $\{\Omega_m^0, H_0, \lambda, 10^4\alpha\}$ for the same data sets as in the previous case. When evaluating the errors from ET standard sirens and comparing them to SNIa + BAO data, we observe that for most parameters, ET’s 1σ constraints are of the same order, apart from the H_0 parameter, which is improved by 1 order of magnitude. This reduction is quantified by the fractional change of $\mathcal{F}_{\Omega_m^0, H_0, \beta, \lambda}^{(\text{SNIa+BAO, ET})} = \{1.1, 0.12, 1.0, 1.1\}$. When the data sets are combined (ET + SNIa + BAO), we find that the 1σ region is narrower for all parameters compared to ET alone, with $\mathcal{F}_{\Omega_m^0, H_0, \alpha, \lambda}^{(\text{ET, ET+SNIa+BAO})} = \{0.84, 0.92, 1.0, 0.92\}$. In the case of LISA standard sirens, we observe that all cosmological and model parameters are better or equally constrained by LISA alone compared to SNIa + BAO, with $\mathcal{F}_{\Omega_m^0, H_0, \alpha, \lambda}^{(\text{SNIa+BAO, LISA})} = \{0.91, 0.13, 1.0, 1.0\}$. Combining LISA with SNIa + BAO, we find improved constraints with respect to the SNIa + BAO data set alone. Moreover, when comparing LISA + SNIa + BAO with LISA alone, the former shows an even better constraining power, with the most significant reduction in error observed for Ω_m^0 , with $\mathcal{F}_{\Omega_m^0, H_0, \alpha, \lambda}^{(\text{LISA, LISA+SNIa+BAO})} = \{0.78, 0.92, 1.0, 0.87\}$. For both the ET and LISA data sets, the accuracy on H_0 can be improved by 1 order of magnitude (0.36 for ET and 0.39 for LISA) compared to SNIa + BAO (3.1), as reported in Sec. IV A as well. Interestingly, for all data sets and combinations, the accuracy of the model parameters remains largely unaffected, with the 1σ region for λ showing only nominal changes and remaining unchanged for α . Comparing the constraining power of ET and LISA with their combination, ET + LISA, we see that the latter provides better constraining power for the cosmological parameters than any of the other data sets analyzed.

TABLE III. Marginalized constraints on cosmological and model parameters for the kinetic model at 68% C.L.

Data sets	Kinetic coupled quintessence							
	Ω_m^0	$\sigma_{\Omega_m^0}$	H_0	σ_{H_0}	$10^4\alpha$	$\sigma_{10^4\alpha}$	λ	σ_λ
SNIa + BAO	$0.3016_{-0.0057}^{+0.0075}$	0.0066	70.4 ± 3.1	3.1	5.1 ± 2.9	2.9	$0.34_{-0.29}^{+0.16}$	0.23
ET	$0.3067_{-0.0046}^{+0.0093}$	0.0070	67.45 ± 0.36	0.36	4.8 ± 2.9	2.9	$0.41_{-0.31}^{+0.20}$	0.26
ET + SNIa + BAO	$0.3062_{-0.0043}^{+0.0074}$	0.0059	67.36 ± 0.33	0.33	5.0 ± 2.9	2.9	$0.37_{-0.28}^{+0.19}$	0.24
LISA	$0.2997_{-0.0041}^{+0.0079}$	0.0060	67.30 ± 0.39	0.39	4.9 ± 2.9	2.9	$0.34_{-0.30}^{+0.16}$	0.23
LISA + SNIa + BAO	$0.3024_{-0.0035}^{+0.0058}$	0.0047	67.47 ± 0.36	0.36	5.0 ± 2.9	2.9	$0.29_{-0.26}^{+0.13}$	0.20
ET + LISA	$0.3040_{-0.0031}^{+0.0058}$	0.0045	67.42 ± 0.26	0.26	5.1 ± 2.9	2.9	$0.31_{-0.26}^{+0.15}$	0.21
ET + LISA + SNIa + BAO	$0.3040_{-0.0031}^{+0.0058}$	0.0045	67.42 ± 0.27	0.27	4.9 ± 2.9	2.9	$0.29_{-0.25}^{+0.14}$	0.20

Regarding the model parameters, there seems to be a minimal change in accuracy compared to the single ET or LISA data sets. We do note that the GW combination provides better accuracy with respect to SNIa + BAO, with $\mathcal{F}_{\Omega_m^0, H_0, \alpha, \lambda}^{(\text{SNIa+BAO, ET+LISA})} = \{0.68, 0.084, 1.0, 0.91\}$. The full combination of ET + LISA + SNIa + BAO, has a negligible change in the constraints when compared to ET + LISA for all parameters.

Comparing the accuracy of the constraints of the kinetic model obtained in Ref. [70] with CMB TT, TE and EE *Planck* 2018, *Planck* CMB lensing, BAO and SNIa data, we note that the parameter α is better constrained by CMB data and its combination with BAO and SNIa by 1 order of magnitude when compared to our data combinations, given that the latter only depend on the background evolution. More precisely, we report $\sigma_{10^4\alpha} = 2.9$ for all the data set combinations while in Ref. [70] this was reduced to $\sigma_{10^4\alpha} = 0.95$ (Plk18), 0.84 (Plk18+SNIa + BAO), and 0.7 (Plk18 + SNIa + BAO + Lensing). Future ET and LISA catalogs will be able to constrain λ at the same level as *Planck* CMB data ($\sigma_\lambda = 0.48$ with Plk18 and $\sigma_\lambda = 0.2$ with both Plk18 + SNIa + BAO and Plk18 + SNIa + BAO + Lensing). On the other hand, the standard siren data will better constrain H_0 by 1 order of magnitude with respect to Plk18 ($\sigma_{H_0} = 2.5$). CMB lensing data increase the constraint by 1 order of magnitude, namely with accuracy $\sigma_{H_0} = 0.6$, which is of the same order of magnitude as the ET and LISA cases, even though the standard sirens perform better in terms of the relative error with $\sigma_{H_0} < 0.4$ for all the combinations considered.

C. Disformal coupling

In the following we study the model with disformal coupling only,

$$C = 1, \quad D = D_0^4 \quad \text{and} \quad V(\phi) = V_0 e^{-\lambda\phi/M_{\text{Pl}}}, \quad (26)$$

in which case the conformal contribution vanishes and D is simply a constant with dimensions of $(\text{mass})^{-4}$ in Eq. (2) [and hence D_0 has units of $(\text{mass})^{-1}$], and $V(\phi)$ follows the same considerations as in the previous cases. The constraints on this model have been obtained in [101,102]. It was found that using background data only [$H(z)$, BAO and supernova Union2.1 data] results in the following constraints: $D_0 > 0.07 \text{ meV}^{-1}$ and $\lambda < 1.56$ at 95.4% [101]. An upper bound can be obtained for D_0 with CMB data (including lensing) and BAO, SNIa, cosmic chronometers, cluster abundance, and H_0 priors which is $D_0 < 0.2500 \text{ meV}^{-1}$ and a stringent upper limit for λ is < 0.6720 at 1σ [102].

From Figs. 8–10, summarized in Table IV, we analyze the results for the parameters $\{H_0, \Omega_m^0, D_0, \lambda\}$ for the same data sets as in the previous cases. In our analysis of the ET data set alone, we observe an improved accuracy for all

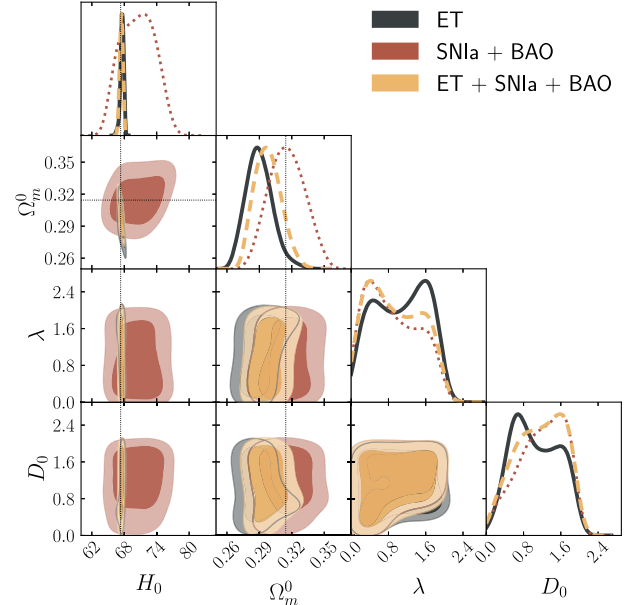


FIG. 8. 68% and 95% C.L. 2D contours and 1D posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, D_0\}$ in the constant disformal coupled quintessence model with ET (charcoal filled line), SNIa + BAO (red dotted line) data and their combination (yellow dashed line). The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

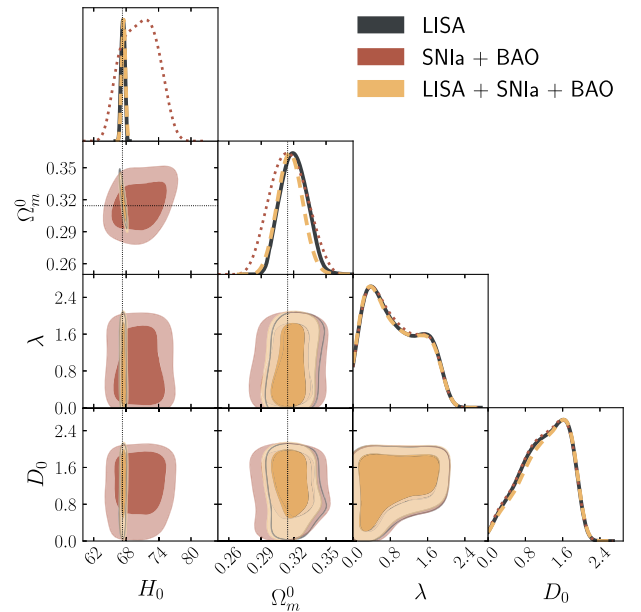


FIG. 9. 68% and 95% C.L. 2D contours and 1D posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, D_0\}$ in the constant disformal coupled quintessence model with LISA (charcoal filled line), SNIa + BAO (red dotted line) data and their combination (yellow dashed line). The scale is the same as in Fig. 8 for comparison purposes, with the SNIa + BAO contours standing as the reference. The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

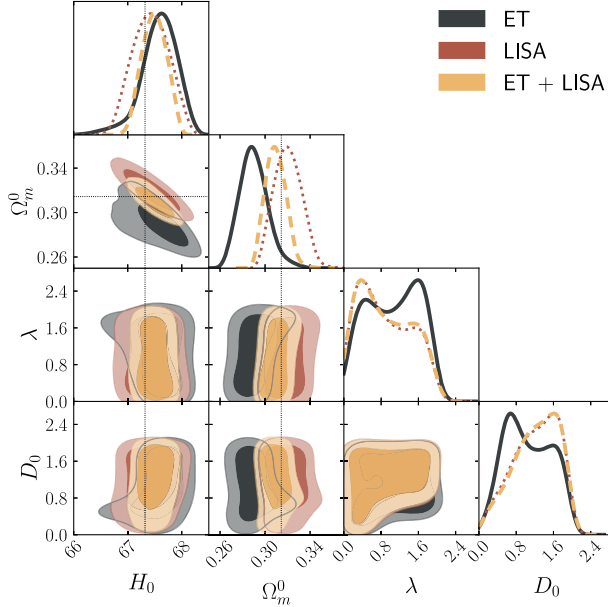


FIG. 10. 68% and 95% C.L. 2D contours and 1D marginalized posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, D_0\}$ in the constant disformal coupled quintessence model with ET mock data (charcoal filled line), LISA mock data (red dotted line) and their combination (yellow dashed line). The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

cosmological and model parameters, $\{\Omega_m^0, H_0, D_0, \lambda\}$, compared to SNIa + BAO, with $\mathcal{F}_{\Omega_m^0, H_0, D_0, \lambda}^{(\text{SNIa+BAO, ET})} = \{0.71, 0.10, 0.98, 0.85\}$. As expected, the combination of data sets, ET + SNIa + BAO, also results in improved accuracy compared to SNIa + BAO. Compared to the ET data set alone, there are only minor changes in the parameters' accuracy, $\mathcal{F}_{\Omega_m^0, H_0, D_0, \lambda}^{(\text{ET, ET+SNIa+BAO})} = \{1.1, 0.97, 1.1, 1.0\}$. In the case of LISA standard sirens, we find that the cosmological and model parameters follow a similar accuracy trend with $\mathcal{F}_{\Omega_m^0, H_0, D_0, \lambda}^{(\text{SNIa+BAO, LISA})} = \{0.71, 0.11, 0.98, 0.85\}$. Moreover, the same is true for the combination LISA + SNIa + BAO,

with $\mathcal{F}_{\Omega_m^0, H_0, D_0, \lambda}^{(\text{LISA, LISA+SNIa+BAO})} = \{1.0, 1.0, 0.98, 1.22\}$. There is only a nominal change in the accuracy of parameters compared to that of LISA alone, apart from λ , which results in a larger 1σ region, with $\sigma_\lambda = 0.71$. Regardless of the data combination, $\{\Omega_m^0, D_0, \lambda\}$ are constrained at the same level, with the parameter λ having a slight improvement in accuracy for both ET and LISA (both have $\sigma_\lambda = 0.58$ compared to $\sigma_\lambda = 0.7$ for SNIa + BAO). The accuracy of the H_0 parameter is 1 order of magnitude better for both ET and LISA than SNIa + BAO.

There is no change in the model parameters for ET and LISA, and thus we see no noticeable change in the constraints for ET + LISA. However, there is an increase in accuracy for the cosmological parameters. As both ET and LISA improved the constraints compared to SNIa + BAO, we observe the expected result, that ET + LISA have further improved constraints with $\mathcal{F}_{\Omega_m^0, H_0, D_0, \lambda}^{(\text{SNIa+BAO, ET+LISA})} = \{0.55, 0.071, 0.96, 0.85\}$. Following the trend with LISA and the combination with SNIa + BAO, we note that the combined data set ET + LISA + SNIa + BAO, has very little change in the accuracy compared to ET + LISA apart from the constraint for λ , which results in a worse accuracy than ET + LISA and SNIa + BAO with $\mathcal{F}_{\Omega_m^0, H_0, D_0, \lambda}^{(\text{ET+LISA, ET+LISA+SNIa+BAO})} = \{1.0, 1.1, 1.0, 1.3\}$.

The first thing to be noted in comparison with the results reported in Ref. [101] is that we were able to derive constraints at 68% C.L. and not only at 95% C.L. for all the model parameters, therefore providing better constraints in all the cases. Moreover, the constraining power on H_0 is largely improved from $\sigma_{H_0} \approx 2.2$ for the background data to $\sigma_{H_0} \approx 0.3$ in all the cases, including standard sirens. When compared with the results of CMB, CMB lensing and additional data in Ref. [102], which reported only upper bounds for λ and D_0 , we see that both parameters are constrained at 68% C.L. with standard sirens, with lower and upper bounds, in particular with more accommodating

TABLE IV. Marginalized constraints on cosmological and model parameters for the constant disformal coupled quintessence model at 68% C.L.

Constant disformal coupled quintessence								
Data sets	Ω_m^0	$\sigma_{\Omega_m^0}$	H_0	σ_{H_0}	D_0/meV^{-1}	σ_{D_0}	λ	σ_λ
SNIa + BAO	0.315 ± 0.017	0.017	70.5 ± 3.1	3.1	$1.20_{-0.38}^{+0.65}$	0.52	$0.87_{-0.76}^{+0.59}$	0.68
ET	$0.290_{-0.013}^{+0.011}$	0.012	$67.58_{-0.27}^{+0.36}$	0.32	1.06 ± 0.51	0.51	1.06 ± 0.58	0.58
ET + SNIa + BAO	$0.298_{-0.014}^{+0.011}$	0.013	67.45 ± 0.31	0.31	$1.15_{-0.44}^{+0.66}$	0.55	0.92 ± 0.58	0.58
LISA	0.320 ± 0.012	0.012	67.43 ± 0.33	0.33	$1.22_{-0.38}^{+0.64}$	0.51	0.87 ± 0.58	0.58
LISA + SNIa + BAO	0.317 ± 0.012	0.012	67.52 ± 0.34	0.34	$1.24_{-0.36}^{+0.64}$	0.50	$0.86_{-0.77}^{+0.65}$	0.71
ET + LISA	$0.3094_{-0.0099}^{+0.0087}$	0.0093	67.49 ± 0.22	0.22	$1.23_{-0.36}^{+0.63}$	0.50	0.88 ± 0.58	0.58
ET + LISA + SNIa + BAO	$0.3100_{-0.0100}^{+0.0092}$	0.0096	67.47 ± 0.25	0.25	$1.24_{-0.36}^{+0.63}$	0.50	$0.88_{-0.77}^{+0.68}$	0.73

upper bounds, as this analysis includes only background data. Accordingly, the error in H_0 is brought to the same order of magnitude with $\sigma_{H_0} \approx 0.9$, which is still about 3 times larger than the one reported in this analysis.

D. Mixed conformal-disformal coupling

Finally, we discuss a model with a mixed coupling consisting of a conformal and a disformal part. Specifically, we consider

$$\begin{aligned} C(\phi) &= e^{2\beta\phi/M_{\text{Pl}}}, & D(\phi) &= D_0^4 \quad \text{and} \\ V(\phi) &= V_0 e^{-\lambda\phi/M_{\text{Pl}}}. \end{aligned} \quad (27)$$

As for the disformal case, the constraints on such a model were discussed in [101,102]. For the same background data it has been reported that $D_0 > 0.102 \text{ meV}^{-1}$, $\beta < 0.453$ and $\lambda < 1.59$ at 95.4% [101]. An example of the constraints including CMB data in [102] are $\beta \lesssim 0.17$ and $\lambda \lesssim 0.35$ at 1σ , with the details depending on the data sets used, with the disformal coupling D_0 not always well constrained for this case, with lower bounds of $D_0 \gtrsim 0.35 \text{ meV}^{-1}$ for some data combinations.

In Figs. 11–13 and Table V we show the results for the parameters $\{H_0, \Omega_m^0, \beta, D_0, \lambda\}$ for the same data sets as before. In our analysis of the ET data set alone, we observe improved accuracy for the cosmological parameters, Ω_m^0 and H_0 , compared to the SNIa + BAO data set, with $\mathcal{F}_{\Omega_m^0, H_0}^{(\text{SNIa+BAO, ET})} = \{0.61, 0.088\}$. The combined data sets,

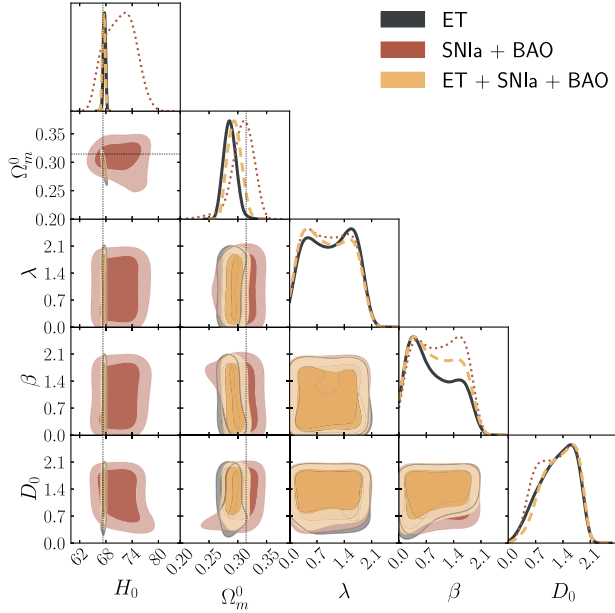


FIG. 11. 68% and 95% C.L. 2D contours and 1D posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, \beta, D_0\}$ in the mixed conformal-disformal coupled quintessence model with ET (charcoal filled line), SNIa + BAO (red dotted line) data and their combination (yellow dashed line). The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

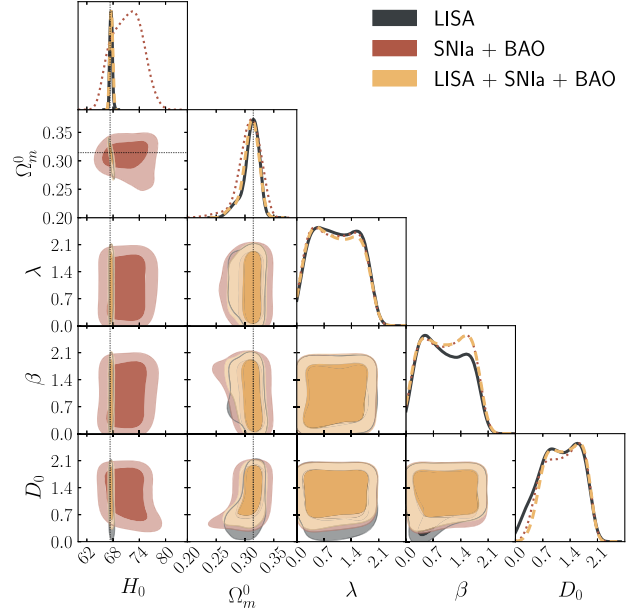


FIG. 12. 68% and 95% C.L. 2D contours and 1D posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, \beta, D_0\}$ in the mixed conformal-disformal coupled quintessence model with LISA (charcoal filled line), SNIa + BAO (red dotted line) data and their combination (yellow dashed line). The scale is the same as in Fig. 11 for comparison purposes, with the SNIa + BAO contours standing as the reference. The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

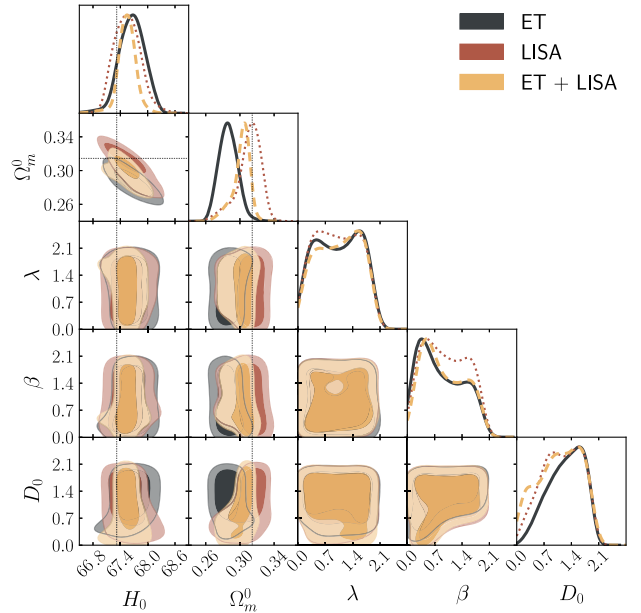


FIG. 13. 68% and 95% C.L. 2D contours and 1D marginalized posterior distributions for the parameters $\{H_0, \Omega_m^0, \lambda, \beta, D_0\}$ in the mixed conformal-disformal coupled quintessence model with ET mock data (charcoal filled line), LISA mock data (red dotted line) and their combination (yellow dashed line). The dotted lines depict the fiducial values for the mock data $\{\Omega_m^0, H_0\} = \{0.3144, 67.32\}$.

TABLE V. Marginalized constraints on cosmological and model parameters for the mixed conformal-disformal coupled quintessence model at 68% C.L.

Mixed conformal-disformal coupled quintessence										
Data sets	Ω_m^0	$\sigma_{\Omega_m^0}$	H_0	σ_{H_0}	β	σ_β	D_0/meV^{-1}	σ_{D_0}	λ	σ_λ
SNIa + BAO	$0.308_{-0.015}^{+0.021}$	0.018	71.2 ± 3.3	3.30	1.01 ± 0.57	0.57	$1.23_{-0.43}^{+0.59}$	0.51	0.98 ± 0.57	0.57
ET	$0.286_{-0.012}^{+0.010}$	0.011	67.65 ± 0.29	0.29	0.85 ± 0.58	0.58	$1.27_{-0.35}^{+0.58}$	0.47	1.03 ± 0.58	0.58
ET + SNIa + BAO	$0.294_{-0.013}^{+0.011}$	0.012	67.50 ± 0.30	0.30	$0.92_{-0.76}^{+0.66}$	0.71	$1.32_{-0.35}^{+0.53}$	0.44	0.97 ± 0.58	0.58
LISA	$0.310_{-0.0087}^{+0.017}$	0.013	$67.55_{-0.31}^{+0.27}$	0.29	0.97 ± 0.56	0.56	$1.15_{-0.44}^{+0.63}$	0.54	1.01 ± 0.56	0.56
LISA + SNIa + BAO	$0.310_{-0.010}^{+0.016}$	0.013	67.59 ± 0.33	0.33	1.01 ± 0.58	0.58	$1.25_{-0.43}^{+0.53}$	0.48	0.98 ± 0.57	0.57
ET + LISA	$0.302_{-0.0058}^{+0.0120}$	0.0089	67.54 ± 0.20	0.20	0.92 ± 0.55	0.55	$1.09_{-0.42}^{+0.76}$	0.59	$1.05_{-0.56}^{+0.71}$	0.64
ET+LISA+SNIa+BAO	$0.304_{-0.0089}^{+0.0120}$	0.0105	67.53 ± 0.24	0.24	0.98 ± 0.57	0.57	$1.27_{-0.41}^{+0.53}$	0.47	0.97 ± 0.57	0.57

ET + SNIa + BAO, show comparable results, with a slight increase in accuracy compared to ET alone. For the model parameters, $\{\beta, D_0, \lambda\}$, ET compared with SNIa + BAO demonstrates close constraining power, with $\mathcal{F}_{\beta, D_0, \lambda}^{(\text{SNIa+BAO,ET})} = \{1.0, 0.92, 1.0\}$. However, the combined data set leads to an increase in the error in β with $\sigma_\beta = 0.71$ for ET + SNIa + BAO. For the case of LISA standard sirens, we find that the cosmological parameters follow a similar trend as ET, with increased accuracy, $\mathcal{F}_{\Omega_m^0, H_0}^{(\text{SNIa+BAO,LISA})} = \{0.72, 0.088\}$, with the combined data set showing a comparable trend. However, it is worth noting a noticeable reduction in accuracy for H_0 between LISA and the combined data set, $\mathcal{F}_{H_0}^{(\text{LISA,LISA+SNIa+BAO})} = \{1.1\}$. Regarding the model parameters, $\{\beta, D_0, \lambda\}$, we find that unlike for ET, LISA alone exhibits increased accuracy compared to SNIa + BAO, except for D_0 , $\mathcal{F}_{\beta, D_0, \lambda}^{(\text{SNIa+BAO,LISA})} = \{0.98, 1.1, 0.98\}$. The combination of the data sets results in comparable accuracy to LISA alone, with $\mathcal{F}_{\beta, D_0, \lambda}^{(\text{LISA,LISA+SNIa+BAO})} = \{1.0, 0.89, 1.0\}$. Similar to Secs. IV A–IV C, the combination of the GW data sets leads to a significant change in the accuracy of Ω_m^0 and H_0 compared to the SNIa + BAO data sets, $\mathcal{F}_{\Omega_m^0, H_0}^{(\text{SNIa+BAO,ET+LISA})} = \{0.49, 0.061\}$. The accuracy of the cosmological parameters is also slightly enhanced compared to ET or LISA alone. Regarding the model parameters, we find only a very small change compared to SNIa + BAO, $\mathcal{F}_{\beta, D_0, \lambda}^{(\text{SNIa+BAO,ET+LISA})} = \{0.96, 1.2, 1.1\}$, noting that both D_0 and λ are slightly less constrained. Combining all of the data sets, we note that there is a similar trend as before, with the cosmological parameters exhibiting an enhanced constraint when compared to SNIa + BAO, while the model parameters remain mostly unchanged, $\mathcal{F}_{\Omega_m^0, H_0, \beta, D_0, \lambda}^{(\text{SNIa+BAO,ET+LISA+SNIa+BAO})} = \{0.58, 0.073, 1.0, 0.92, 1.0\}$.

In summary, regardless of the combination of data sets considered, the constraints on $\Omega_m^0, \beta, D_0, \lambda$ are of the same order of magnitude as those obtained from SNIa + BAO. Additionally, we note that the accuracy on H_0 is improved by 1 order of magnitude for both ET and LISA compared to SNIa + BAO.

Similarly to the comparison in Sec. IV C, the main improvement in contrast with the results reported in Ref. [101] is the fact that we can obtain constraints at 68% C.L. for all the model parameters. The potential parameter λ is constrained with upper bounds for the standard sirens at 1σ of the same order of the 2σ ones reported in the previous studies. Moreover, the constraining power on H_0 is largely improved from $\sigma_{H_0} \approx 2.1$ for the background data to $\sigma_{H_0} \approx 0.3$ in all the cases, including standard sirens. The comparison with results including CMB, CMB lensing and additional data in Ref. [102], which are either unable to constrain D_0 or find just a lower bound and report only upper bounds for λ and β , shows that standard sirens successfully constrain the three model parameters at 68% C.L. for all the combinations, which is a great improvement given that only background data has been considered. Including CMB data brings the error in H_0 to the same order of magnitude with $\sigma_{H_0} \approx 0.6$, which is still around 2 times larger than the ones reported in this analysis.

V. CONCLUSIONS

In this paper, we have explored the potential of future GWs detectors, namely LISA and ET, to constrain conformal and disformal couplings between the dark energy and dark matter fluids. We have considered four models: conformal coupled quintessence, a kinetic model, constant disformal coupled quintessence and a mixed conformal-disformal model. All the cases considered have the same exponential potential of which we have constrained its slope, λ .

We have generated mock catalogs of standard siren events with ET and LISA specifications and with those we performed an MCMC analysis considering, separately, their combination with current SNIa + BAO data as well for reference. Under the assumption we have used to generate the mock data assuming a particular cosmology, we find the following:

- (1) The conformal coupled quintessence: the combinations of LISA + SNIa + BAO and ET + SNIa + BAO improve the constraints on both the slope parameter, λ , and the conformal coupling parameter β . The combination of ET + LISA with SNIa + BAO reduces the error in β by one third.
- (2) The kinetic model: ET and LISA alone cannot improve the constraints on λ or on the conformal exponential parameter α . When LISA is combined with SNIa + BAO the accuracy slightly improves for the slope parameter and for the matter density Ω_m^0 .
- (3) The constant disformal coupled quintessence: all combinations can constrain the disformal parameter D_0 at 1σ with the same order of magnitude and a small improvement for LISA + SNIa + BAO. For the slope parameter instead, both ET, LISA and their combination perform better than SNIa + BAO. Moreover for the full catalog combination of ET + LISA, the error in Ω_m^0 can be reduced.
- (4) The mixed conformal-disformal coupled quintessence: for all the parameters of the model there is not a significant improvement on the accuracy of their constraints when using ET or LISA data separately. A small reduction in the size of the 1σ region is reported only for the disformal parameter, D_0 , in the

full combinations. The error in Ω_m^0 is slightly reduced for ET + LISA.

Regardless of the model considered we found that the accuracy on the H_0 parameter increases by 1 order of magnitude at 1σ when compared to the combination of BAO and SNIa data. This is promising in light of solving/understanding the H_0 tension. This improvement is also responsible for the increased accuracy on the constraints for the model parameters when we consider the full combinations that we just reviewed.

Ultimately, our results show that future 3G detectors can improve our knowledge on DE-DM interaction and shed light on the H_0 tension.

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