



# THE INTERMEDIATE LUMINOSITY OPTICAL TRANSIENT SN 2010DA: THE PROGENITOR, ERUPTION, AND AFTERMATH OF A PECULIAR SUPERGIANT HIGH-MASS X-RAY BINARY

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## ABSTRACT

We present optical spectroscopy, ultraviolet-to-infrared imaging, and X-ray observations of the intermediate luminosity optical transient (ILOT) SN 2010da in NGC 300 ( $d = 1.86$  Mpc) spanning from  $-6$  to  $+6$  years relative to the time of outburst in 2010. Based on the light-curve and multi-epoch spectral energy distributions of SN 2010da, we conclude that the progenitor of SN 2010da is a  $\approx 10$ – $12 M_{\odot}$  yellow supergiant possibly transitioning into a blue-loop phase. During outburst, SN 2010da had a peak absolute magnitude of  $M_{\text{bol}} \lesssim -10.4$  mag, dimmer than other ILOTs and supernova impostors. We detect multi-component hydrogen Balmer, Paschen, and Ca II emission lines in our high-resolution spectra, which indicate a dusty and complex circumstellar environment. Since the 2010 eruption, the star has brightened by a factor of  $\approx 5$  and remains highly variable in the optical. Furthermore, we detect SN 2010da in archival *Swift* and *Chandra* observations as an ultraluminous X-ray source ( $L_X \approx 6 \times 10^{39}$  erg s<sup>-1</sup>). We additionally attribute He II 4686 Å and coronal Fe emission lines in addition to a steady X-ray luminosity of  $\approx 10^{37}$  erg s<sup>-1</sup> to the presence of a compact companion.

*Key words:* stars: mass-loss – supernovae: individual (SN 2010da) – X-rays: binaries

## 1. INTRODUCTION

Between the luminosities of the brightest novae ( $M_V \approx -10$ ; Hachisu & Kato 2014) and the dimmest supernovae ( $M_V \approx -14$ ; Zampieri et al. 2003), there is a dearth of well-studied optical transients (see Kasliwal 2012). In the last decade, we have begun to fill in this gap with a number of exotic events such as luminous red novae (Kulkarni et al. 2007), luminous blue variable (LBV) outbursts, and other “supernova impostors” (e.g., Van Dyk et al. 2000; Pastorello et al. 2007; Berger et al. 2009; Tartaglia et al. 2015). Additionally, there are expected events that have not been definitively observed, such as “failed” supernovae (Kochanek et al. 2008). Following Berger et al. (2009), we will collectively refer to these events as intermediate luminosity optical transients (ILOTs).

The link between ILOTs and their progenitors remains elusive, especially for ILOTs surrounded by dense circumstellar media (CSM). Brighter dusty ILOTs, such as the great eruption of Eta Carinae (Davidson & Humphreys 1997) or SN 1954J (Van Dyk et al. 2005), are attributed to LBV outbursts; however, the progenitors of dimmer events are under debate with a larger pool of viable origins. For example, theorized progenitors of the famous dusty ILOTs, such as NGC 300 OT2008-1 and SN 2008S have ranged from mass-loss events of yellow hypergiants (Berger et al. 2009), to mass transfer from an extreme AGB star to a main-sequence companion (Kashi et al. 2010), to low luminosity electron-capture supernovae (Thompson et al. 2009; Adams et al. 2016). Each of these

interpretations shares the common theme of marking an important point in the evolution of relatively massive stars ( $\gtrsim 9 M_{\odot}$ ).

Adding to the diversity of ILOTs is the possibility of optical transients within X-ray binary systems. High-mass X-ray binaries (HMXBs) consist of a massive star and a compact object (e.g., a neutron star or a black hole) and produce X-rays as material accretes onto the compact object through a variety of channels (Lewin et al. 1997; Reig 2011). A relatively new subclass of HMXBs, known as obscured HMXBs, are cloaked in a high density of local material ( $N_H \sim 10^{23}$ – $10^{24}$  cm<sup>-2</sup>; Chaty & Rahoui 2007; Tomsick et al. 2009). While the primary stars of these systems are largely unknown, several have been shown to be supergiants exhibiting B[e] phenomena (Clark et al. 1999; Chaty & Filliatre 2005; Kaplan et al. 2006). These systems likely produce their dense circumstellar material through either a continuous wind or periodic outbursts, which have not yet been observed.

In this work, we report data from a five-year, multi-wavelength (X-ray, ultraviolet, optical, and infrared) observational campaign of the dusty ILOT SN 2010da, which was discovered in the nearby galaxy NGC 300 (Monard 2010). We show that SN 2010da exhibits many features shared among dusty ILOTs, such as striking Balmer emission and optical variability on the order of months, but it is the only ILOT to sit in an intermediate range between extremely dusty red transients such as SN 2008S and the bluer, brighter LBV outbursts. Additionally, SN 2010da is the first ILOT to be a member of an HMXB, which undergoes an ultraluminous X-ray outburst ( $\sim 10^{40}$  erg s<sup>-1</sup>). Previous work on SN 2010da (Binder et al. 2011, 2016) concluded that the progenitor is a massive

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( $\gtrsim 25 M_{\odot}$ ) LBV using limited *Hubble Space Telescope* (*HST*) photometry. However, from our broadband photometry and spectroscopy, we infer that SN 2010da originated from an intermediate mass ( $\sim 10\text{--}12 M_{\odot}$ ), variable yellow supergiant progenitor, which is now transitioning into a blue-loop phase of its evolution. We discuss these conflicting interpretations and the importance of comprehensive, multiwavelength coverage of ILOTs.

## 2. OBSERVATIONS

SN 2010da was discovered in NGC 300 on 2010 May 23.169 UT by Monard (2010) with an unfiltered magnitude of  $16.0 \pm 0.2$ , corresponding to  $M \approx -10.3$  assuming a distance of 1.86 Mpc (Rizzi et al. 2006) and a foreground extinction of  $E(B-V) = 0.011$  (Schlafly & Finkbeiner 2011). We neglect additional extinction from NGC 300 based on our observed *Swift* colors (Section 3.1.2). Throughout this paper, Epoch 0 will refer to the discovery date, 2010 May 23. Prior to discovery, NGC 300 was behind the Sun, though Monard (2010) reported an upper limit of  $\lesssim 15.5$  mag on May 6. Archival *Spitzer* data indicated that the source began brightening in the infrared at least 150 days before the optical discovery (Laskar et al. 2010). Multiwavelength follow-up, spanning from the radio to X-rays, revealed that despite its supernova designation, SN 2010da was likely an outburst of a massive star enshrouded by dust (Chornock & Berger 2010; Elias-Rosa et al. 2010; Prieto et al. 2010). This conclusion was reaffirmed by archival *Spitzer*/IRAC observations of the dusty progenitor (Berger & Chornock 2010; Khan et al. 2010), but the lack of extinction in the spectral energy distribution (SED) suggested that some dust had been destroyed during the outburst (Bond 2010; Brown 2010). Early spectroscopic follow-up revealed narrow emission features (FWHM  $\approx 1000 \text{ km s}^{-1}$ ) with no signs of P-Cygni profiles (Elias-Rosa et al. 2010). Hydrogen Balmer, Fe II, and He I emission lines provided further support for interaction with a dense CSM surrounding the progenitor.

The transient was also detected in the X-rays and UV with the *Swift* X-ray Telescope (XRT) and Ultraviolet/Optical Telescope (UVOT), respectively (Brown 2010; Immler et al. 2010). Additionally,  $3\sigma$  upper limits of  $F_{\nu} \lesssim 87$  (4.9 GHz),  $\lesssim 75$  (8.5 GHz), and  $\lesssim 225$  (22.5 GHz)  $\mu\text{Jy}$  were obtained with the Karl G. Jansky Very Large Array (Chomiuk & Soderberg 2010). Following the event, we monitored SN 2010da in the near-infrared (NIR) and optical using Gemini and *Magellan*. We report below our ground-based imaging and spectroscopy, as well as an analysis of archival *Spitzer*, *Hubble*, *Swift*, and *Chandra* observations.

### 2.1. *Spitzer* Infrared Imaging

We obtained publicly available *Spitzer* images spanning from 2003 November 21 to 2016 March 19 (see Table 1 for program IDs; Lau et al. 2016). This data set extends several years before and after the event, but no data are available within a four month window surrounding the optical discovery. We used data from the InfraRed Array Camera (IRAC) in the 3.6 and 4.5  $\mu\text{m}$  bands through both the original and “warm” *Spitzer* missions, and we use IRAC data in the 5.8 and 8.0  $\mu\text{m}$  bands available prior to the 2010 eruption. Additionally, we used photometry from the Multiband Imagine Photometer (MIPS) in the 24  $\mu\text{m}$  band prior to the discovery of the transient. We processed the *Spitzer* data with the *Mopex*

package, which creates a mosaic of the dithered *Spitzer* images. For the IRAC images, we used a drizzling parameter of 0.7 and an output pixel scale of  $0''.4$ . For the MIPS images, we used the same drizzling parameter but with an output pixel scale of  $1''.8$ . Images of the field in the *Spitzer* bands are shown in Figure 1.

We performed aperture photometry using DS9’s *Funtools*. For the IRAC photometry, we used an aperture of three native IRAC pixels (corresponding to  $3''.66$ ) with inner and outer background annulus radii of 3 ( $3''.66$ ) and 7 ( $8''.54$ ) native pixels, respectively. These radii have calculated aperture correction factors for point sources in the IRAC Instrument Handbook. For the MIPS 24  $\mu\text{m}$  photometry, we used an aperture size of  $3''.5$  with no background annulus, following the MIPS Instrument Handbook. We calculated the flux uncertainties following Equation 1 in Laskar et al. (2011). The observations are summarized in Table 1, and the *Spitzer*/IRAC light curves at 3.6 and 4.5  $\mu\text{m}$  are shown in Figure 2. Our photometric results are consistent with those presented in Lau et al. (2016).

### 2.2. Ground-based NIR Imaging

We obtained NIR imaging observations with the FourStar Infrared Camera (Persson et al. 2013) on the *Magellan*/Baade 6.5 m telescope at the Las Campanas Observatory in Chile on three epochs: 2011 December 7 (*J, H, K<sub>s</sub>*), 2015 July 31 (*H, K<sub>s</sub>*), and 2015 August 18 (*J, H, K<sub>s</sub>*). We calibrated, aligned, and co-added each of these observations using the FSRED package.<sup>10</sup> Each image was calibrated using the 2MASS Point Source Catalog, and the magnitude of the transient was measured using aperture photometry. The results are summarized in Table 2.

### 2.3. Ground-based Optical Imaging

We obtained optical imaging observations with the Low Dispersion Survey Spectrograph 3 (LDSS-3, upgraded from LDSS-2 Allington-Smith et al. 1994) and the Inamori-*Magellan* Areal Camera and Spectrograph (IMACS; Dressler et al. 2006) on the *Magellan* Clay and Baade 6.5 m telescopes at the Las Campanas Observatory, respectively, in the *gri* filters spanning from  $\approx 610$  days before to  $\approx 1900$  days after the optical discovery. In our earliest IMACS I-band image (at Epoch  $-609$ ), we detect the object with  $24.2 \pm 0.2$  mag (see Figure 1). However, we do not detect a source at the location of SN 2010da in pre-transient *gri* images taken with the *Magellan*/Clay wide field imager MegaCam (at Epoch  $-183$ ; McLeod et al. 2015). We use the MegaCam images in each filter as templates for image subtraction. For all other ground-based optical imaging observations, we performed image subtraction using the ISIS package (Alard 2000). We then performed aperture photometry on the subtracted images and calibrated to southern standard stars listed in Smith et al. (2007). The photometry is summarized in Table 3.

### 2.4. *HST* Optical Imaging

SN 2010da was observed by the *HST* Advanced Camera for Surveys (ACS) on 2012 July 18 (Program 12450) and 2014 July 9 (Program 13515). The object was observed in the F814W filter in both programs (2224 s and 2548 s exposure

<sup>10</sup> <http://instrumentation.obs.carnegiescience.edu/FourStar/SOFTWARE/reduction.html>

**Table 1**  
*Spitzer* Photometry

Instrument	AOR	PI	Date (UT)	Epoch (Days)	Filter	AB Magnitude
<i>IRAC</i>	6069760	Helou	2003 Nov 21	-2375	3.6	18.77 ± 0.10
<i>IRAC</i>	6069760	Helou	2003 Nov 21	-2375	4.5	18.55 ± 0.07
<i>IRAC</i>	6069760	Helou	2003 Nov 21	-2375	5.8	19.10 ± 0.52
<i>IRAC</i>	6069760	Helou	2003 Nov 21	-2375	8	19.51 ± 0.76
<i>MIPS</i>	22611456	Kennicutt	2007 Jul 16	-1042	24	>17.00
<i>IRAC</i>	22517504	Kennicutt	2007 Dec 29	-876	3.6	18.79 ± 0.07
<i>IRAC</i>	22517504	Kennicutt	2007 Dec 29	-876	4.5	18.67 ± 0.05
<i>IRAC</i>	22517504	Kennicutt	2007 Dec 29	-876	5.8	19.50 ± 0.51
<i>IRAC</i>	22517504	Kennicutt	2007 Dec 29	-876	8	>17.45
<i>IRAC</i>	31527680	Freedman	2009 Dec 21	-153	3.6	18.39 ± 0.07
<i>IRAC</i>	31527424	Freedman	2010 Jan 13	-130	3.6	17.84 ± 0.04
<i>IRAC</i>	31528448	Freedman	2010 Jul 27	65	3.6	17.87 ± 0.04
<i>IRAC</i>	31528192	Freedman	2010 Aug 16	85	3.6	18.11 ± 0.05
<i>IRAC</i>	31527936	Freedman	2010 Aug 31	100	3.6	18.36 ± 0.07
<i>IRAC</i>	42195968	Kochanek	2011 Aug 29	463	3.6	18.68 ± 0.09
<i>IRAC</i>	42195968	Kochanek	2011 Aug 29	463	4.5	18.85 ± 0.08
<i>IRAC</i>	42502912	Kasliwal	2012 Jan 14	601	3.6	18.66 ± 0.08
<i>IRAC</i>	42195712	Kochanek	2012 Aug 10	810	3.6	18.41 ± 0.07
<i>IRAC</i>	42195712	Kochanek	2012 Aug 10	810	4.5	18.58 ± 0.07
<i>IRAC</i>	50572032	Kasliwal	2014 Mar 13	1390	3.6	18.41 ± 0.07
<i>IRAC</i>	50572032	Kasliwal	2014 Mar 13	1390	4.5	18.65 ± 0.07
<i>IRAC</i>	50573056	Kasliwal	2014 Sep 05	1566	3.6	18.16 ± 0.05
<i>IRAC</i>	50573056	Kasliwal	2014 Sep 05	1566	4.5	18.23 ± 0.05
<i>IRAC</i>	50572544	Kasliwal	2014 Oct 03	1594	3.6	18.21 ± 0.06
<i>IRAC</i>	50572544	Kasliwal	2014 Oct 03	1594	4.5	18.28 ± 0.05
<i>IRAC</i>	50044672	Fox	2014 Oct 14	1605	3.6	18.26 ± 0.06
<i>IRAC</i>	50044672	Fox	2014 Oct 14	1605	4.5	18.34 ± 0.04
<i>IRAC</i>	53022208	Kochanek	2015 Feb 09	1723	3.6	18.33 ± 0.06
<i>IRAC</i>	52691712	Kasliwal	2015 Sep 22	1948	3.6	17.91 ± 0.05
<i>IRAC</i>	52691712	Kasliwal	2015 Sep 22	1948	4.5	18.03 ± 0.03
<i>IRAC</i>	52691968	Kasliwal	2015 Sep 29	1955	3.6	17.90 ± 0.04
<i>IRAC</i>	52691968	Kasliwal	2015 Sep 29	1955	4.5	18.03 ± 0.03
<i>IRAC</i>	52692224	Kasliwal	2015 Oct 12	1968	3.6	17.89 ± 0.05
<i>IRAC</i>	52692224	Kasliwal	2015 Oct 12	1968	4.5	17.99 ± 0.03
<i>IRAC</i>	52692480	Kasliwal	2016 Feb 22	2101	3.6	18.09 ± 0.06
<i>IRAC</i>	52692480	Kasliwal	2016 Feb 22	2101	4.5	18.18 ± 0.04
<i>IRAC</i>	52692736	Kasliwal	2016 Feb 29	2108	3.6	18.13 ± 0.05
<i>IRAC</i>	52692736	Kasliwal	2016 Feb 29	2108	4.5	18.19 ± 0.04
<i>IRAC</i>	52692992	Kasliwal	2016 Mar 19	2127	3.6	18.23 ± 0.06
<i>IRAC</i>	52692992	Kasliwal	2016 Mar 19	2127	4.5	18.22 ± 0.05

times, respectively) and in the F606W filter with program 13515 (2400 s). We processed the data using the standard `PyDrizzle` pipeline in `PyRAF`, which supplies geometric distortion corrections to combine undersampled, dithered images from *HST*. We scaled the pixel size by 0.8 for a final pixel scale of  $0''.032$ . We detected a source coincident with the position of SN 2010da, and using five objects detected in the field from the 2MASS Point Source Catalog, we determined a position of  $\alpha = 00^{\text{h}}55^{\text{m}}04.86^{\text{s}}$ ,  $\delta = -37^{\circ}41'43''.8$  (J2000) with  $0''.3$  uncertainty in both coordinates. This is in good agreement (within  $1\sigma$ ) with previous results (Monard 2010; Binder et al. 2011). With the high resolution of *HST*, SN 2010da appears isolated, and we used aperture photometry to measure its magnitude. These magnitudes are listed in Table 4 and are in good agreement with those reported by Binder et al. (2016).

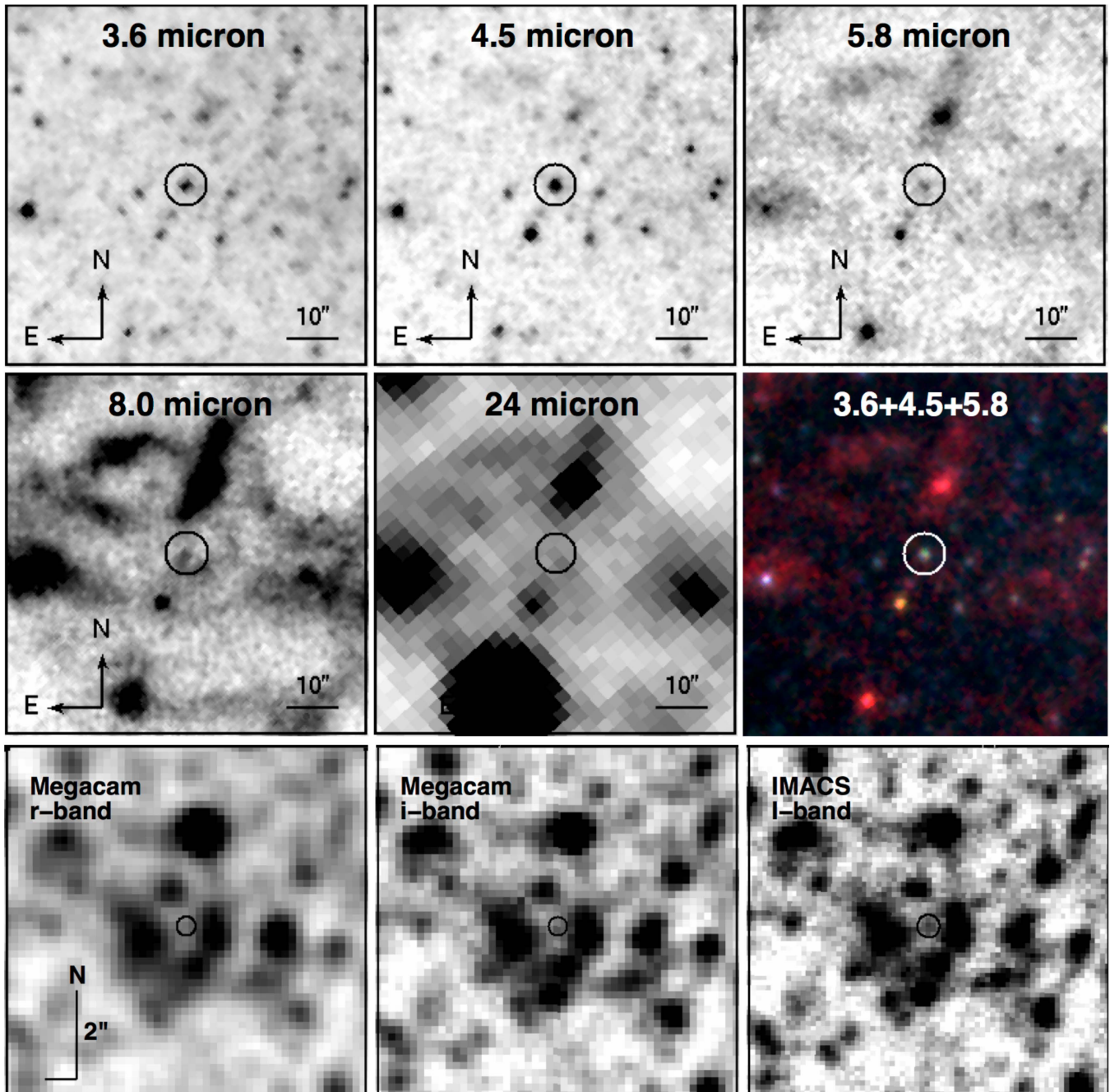
### 2.5. Optical Spectroscopy

We obtained medium- and high-resolution spectra of SN 2010da using the Gemini South Multi-Object Spectrograph (GMOS; Davies et al. 1997), located in the

southern Gemini Observatory in Chile; IMACS, the *Magellan* Inamori Kyocera Echelle (MIKE; Bernstein et al. 2003) spectrograph on the 6.5 m *Magellan*/Clay telescope; and the *Magellan* Echellette Spectrograph (MagE; Marshall et al. 2008) also mounted on the *Magellan*/Clay telescope. Table 5 summarizes these observations. We used standard `IRAF` routines to process the spectra and applied wavelength calibrations using HeNeAr arc lamps. The MIKE spectra were processed using a custom pipeline and calibrated using ThAr arc lamps. We used our own IDL routines to apply flux calibrations from observations of standard stars (archival in the case of Gemini) and correct for telluric absorption. We estimate the resolution of each spectrum (see Table 5) using their associated arc lamp spectra. All spectra are corrected for air-to-vacuum and heliocentric shifts.

### 2.6. Swift/UVOT Imaging

The *Swift*/UVOT data was processed using the method of the *Swift* Optical/Ultraviolet Supernova Archive (SOUSA; Brown et al. 2014). We combined pre-outburst observations



**Figure 1.** Top rows: *Spitzer* images of the SN 2010da progenitor. The right panel of the middle row shows a false color image combining the 3.6 (blue), 4.5 (green), and 5.8 (red)  $\mu\text{m}$  images. Bottom row: Archival MegaCam and IMACS images. The progenitor is only detected in the IMACS I-band image.

from 2006 December and 2007 January into templates from which the underlying host galaxy count rate was measured. A  $3''$  aperture was used with aperture corrections based on an average PSF. A time-dependent sensitivity correction was used (updated in 2015) and AB zeropoints from Breeveld et al. (2011). The photometry is summarized in Table 6.

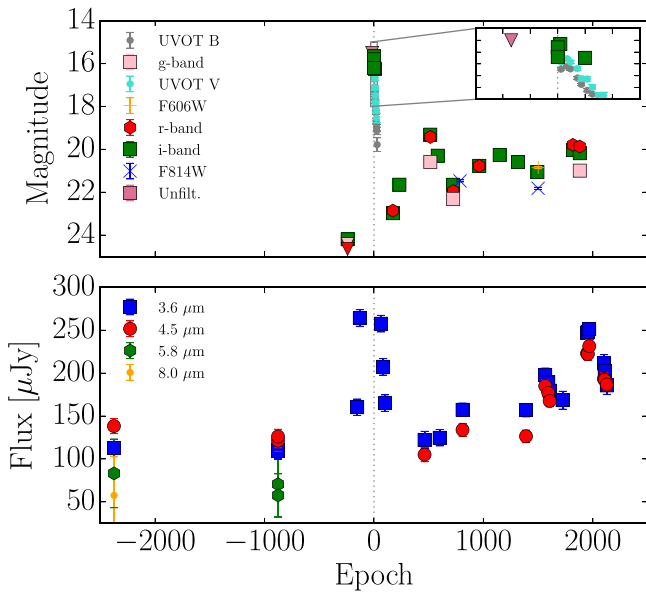
### 2.7. X-Ray Spectral Imaging

We aggregated archival X-ray observations from the *Swift*/XRT, the *Chandra* X-ray Observatory, and *XMM-Newton*. These X-ray observations span from 2000 December 26 to 2014 November 17, including the outburst period. The source

was undetected with *XMM-Newton*, and we use the  $3\sigma$  upper limits obtained by Binder et al. (2011).

The XRT observations were made before, during, and after the 2010 outburst, and an X-ray source coincident with SN 2010da is detected in all three regimes. These observations are publicly available from the *Swift* Archive (Evans et al. 2009), and the XRT photometry and spectra are automatically generated through this database.

We used three archival *Chandra* observations from 2010 September 24 (Obs. ID: 12238; PI: Williams), 2014 June 16 (Obs. ID: 16028; PI: Binder), and 2014 November 17 (Obs. ID: 16029; PI: Binder). All observations were made using the Advanced CCD Imaging Spectrometer (ACIS-I) with similar



**Figure 2.** *Spitzer*, ground-based optical, *HST*, and *UVOT* light curves of SN 2010da, spanning six years before to six years after the 2010 eruption. Epoch 0 refers to 2010 May 23.169 UT, the date of discovery. Downward facing triangles are  $3\sigma$  upper limits.

**Table 2**  
Magellan/FourStar Photometry

Date (UT)	Epoch	Filter	AB Magnitude
2011 Dec 07	563	<i>J</i>	$14.47 \pm 0.09$
2011 Dec 07	563	<i>H</i>	$14.17 \pm 0.11$
2011 Dec 07	563	<i>K<sub>s</sub></i>	$14.21 \pm 0.12$
2015 Jul 31	1895	<i>J</i>	$14.23 \pm 0.02$
2015 Jul 31	1895	<i>H</i>	$13.76 \pm 0.03$
2015 Jul 31	1895	<i>K<sub>s</sub></i>	$13.57 \pm 0.02$
2015 Aug 18	1913	<i>H</i>	$13.97 \pm 0.01$
2015 Aug 18	1913	<i>K<sub>s</sub></i>	$13.63 \pm 0.02$

exposure times (63.0 ks, 64.24 ks, and 61.27 ks, respectively). We analyzed the observations using CIAO version 4.7 and CALDB version 4.6.7 using standard extraction procedures. We performed photometry with WAVDETECT using an annular background region with an inner radius of  $24''.6$  and a width of  $4''.9$  centered on the source. The results are summarized in Table 7. We extracted spectra of the source using the built-in function `specextract`.

### 3. THE MULTIWAVELENGTH PROPERTIES OF SN 2010DA, ITS PROGENITOR, AND ITS PROGENY

#### 3.1. Light-curve and Spectral Evolution

##### 3.1.1. The Progenitor

We are able to constrain the progenitor properties using the *Spitzer* (3.6, 4.5, 5.8, 8, and  $24\ \mu\text{m}$ ) and MegaCam/IMACS (*gri*) observations. We note that the location of SN 2010da was observed in the *i'*-band on both 2008 September 09 and 2009 November 25 by IMACS and MegaCam, respectively. The MegaCam/IMACS observations are summarized in Table 3. The progenitor IMACS detection and MegaCam upper limit are consistent with a magnitude  $\sim 24.2$ . The *gr* upper limits were both obtained with MegaCam on 2009 November 25. The location of SN 2010da was observed five times by *Spitzer*

**Table 3**  
Ground-based Optical Photometry

Date (UT)	Epoch (days)	Instrument	Filter	AB Magnitude
2008 Sep 09	-609	IMACS	<i>i'</i>	$24.19 \pm 0.20$
2009 Nov 25	-179	MegaCam	<i>i'</i>	$>24.4$
2009 Nov 25	-179	MegaCam	<i>r'</i>	$>24.4$
2009 Nov 25	-179	MegaCam	<i>g'</i>	$>24.4$
2010 Nov 13	174	IMACS	<i>i'</i>	$22.97 \pm 0.06$
2010 Nov 13	174	IMACS	<i>r'</i>	$22.85 \pm 0.04$
2011 Jan 12	234	LDSS-3	<i>i'</i>	$21.64 \pm 0.03$
2011 Oct 21	516	LDSS-3	<i>i'</i>	$19.77 \pm 0.06$
2011 Oct 21	516	LDSS-3	<i>r'</i>	$19.42 \pm 0.04$
2011 Oct 21	516	LDSS-3	<i>g'</i>	$20.58 \pm 0.03$
2011 Dec 27	583	IMACS	<i>i'</i>	$20.29 \pm 0.05$
2011 Dec 27	583	IMACS	<i>r'</i>	$20.27 \pm 0.09$
2011 Dec 27	583	IMACS	<i>g'</i>	$22.29 \pm 0.20$
2012 May 17	725	LDSS-3	<i>i'</i>	$21.65 \pm 0.04$
2012 May 17	725	LDSS-3	<i>r'</i>	$21.94 \pm 0.04$
2012 May 17	725	LDSS-3	<i>g'</i>	$22.32 \pm 0.03$
2013 Jan 11	964	LDSS-3	<i>i'</i>	$21.70 \pm 0.04$
2013 Jan 11	964	LDSS-3	<i>r'</i>	$20.75 \pm 0.06$
2013 Jul 15	1149	LDSS-3	<i>i'</i>	$20.06 \pm 0.01$
2013 Dec 30	1317	LDSS-3	<i>i'</i>	$20.80 \pm 0.18$
2014 Jun 26	1495	LDSS-3	<i>i'</i>	$21.03 \pm 0.01$
2015 May 15	1818	IMACS	<i>i'</i>	$20.03 \pm 0.02$
2015 May 15	1818	IMACS	<i>r'</i>	$18.89 \pm 0.13$
2015 Jul 17	1881	IMACS	<i>i'</i>	$20.18 \pm 0.02$
2015 Jul 17	1881	IMACS	<i>r'</i>	$19.87 \pm 0.02$
2015 Jul 17	1881	IMACS	<i>g'</i>	$20.99 \pm 0.03$
2015 Aug 01	1896	IMACS	<i>i'</i>	$19.71 \pm 0.05$
2015 Aug 01	1896	IMACS	<i>r'</i>	$20.09 \pm 0.05$
2015 Aug 01	1896	IMACS	<i>g'</i>	$20.74 \pm 0.08$

before the transient, ranging between 2003 November 21 and 2010 January 14. These observations are summarized in Table 1. We find no significant change in the color and brightness between the pre-eruption observations.

To create a progenitor SED, we average the two pre-eruption *Spitzer* observations in the 3.6 and  $4.5\ \mu\text{m}$  filters and compiled the other detections. The SED of the progenitor is well fit by an unabsorbed blackbody spectrum with  $T = 1500 \pm 40\ \text{K}$  and  $R = 9.4 \pm 0.5\ \text{au}$  ( $\chi_r^2 = 1.2$  for d.f. = 3). These parameters correspond to a bolometric luminosity of  $L = (1.92 \pm 0.26) \times 10^4 L_\odot$ , suggesting a  $\sim 15 M_\odot$  main-sequence progenitor if we assume solar metallicity (Meynet & Maeder 2000). The large radius and cool temperature of this fit imply that the progenitor is surrounded by dust. The progenitor SED is shown in Figure 3 along with several red supergiants (RSGs) and the progenitor of a previous ILOT in NGC 300 (NGC 300 OT2008-1; Berger et al. 2009). Also shown is the SED of an obscured HMXB (IGR J16207-5129; Tomsick et al. 2006). The progenitor SED peaks between the typically bluer obscured HMXBs and the redder ILOTs such as NGC 300 OT2008-1. The SEDs of RSGs seem to bridge this gap, owing their SED variability to diverse geometries (e.g., WOH G64 has notable IR excess possibly due to a dusty torus along the line of sight; Ohnaka et al. 2008), though neither RSG fits the observed SED.

##### 3.1.2. The 2010 Outburst

SN 2010da was discovered at its brightest known magnitude of  $m_{\text{unfiltered}} = 16.0 \pm 0.2$ . It is unclear if SN 2010da was caught at its true peak brightness, but the optical 15.5 mag

**Table 4**  
HST Photometry

Start Date (UT)	Epoch	Proposal ID	PI	Filter	AB Magnitude
2012 Jul 18	787	12450	Kochanek	F814W	20.63 ± 0.03
2014 Jul 02	1501	13515	Binder	F606W	20.68 ± 0.02
2014 Jul 02	1501	13515	Binder	F814W	20.99 ± 0.03

upper limit 18 days prior and a slight rise in the *Swift*/UVOT light curve hints that SN 2010da was discovered near its peak luminosity. An increase in IR flux is seen in the IRAC data as early as  $\approx 130$  days before the optical discovery. The full rise and fall caught by *Spitzer* spans  $\approx 250$  days, as shown in Figure 2.

During the 2010 outburst, the SED of SN 2010da is well fit by two unabsorbed blackbodies at  $\sim 0.2\text{--}1.7\ \mu\text{m}$ : a hotter blackbody with  $T_{H,1} = 9440 \pm 280\ \text{K}$  and  $R_{H,1} = 1.59 \pm 0.14\ \text{au}$  and a cooler blackbody with  $T_{C,1} = 3230 \pm 490\ \text{K}$  and  $R_{C,1} = 9.5 \pm 2.9\ \text{au}$  ( $\chi_r^2 \approx 0.8$  for d.f. = 6), as shown in Figure 4. These blackbodies have a combined bolometric luminosity of  $L = (1.3 \pm 0.4) \times 10^6 L_\odot$ , about 60 times more luminous than the progenitor. Nine days later, the SED is consistent with similar blackbodies, though the larger one has cooled ( $T_{C,2} = 2760 \pm 250\ \text{K}$ ,  $R_{C,2} = 10.5 \pm 1.6\ \text{au}$ ), while the hotter one remains at roughly the same temperature ( $T_{H,2} = 9080 \pm 330\ \text{K}$ ,  $R_{H,2} = 1.25 \pm 0.13\ \text{au}$ ). The radius of the cooler, larger blackbody component is consistent with the estimated pre-eruption radius ( $R_{C,2} \approx R_{C,1} \approx R_{C,0}$ ) but has a temperature that is twice as high ( $T_{C,2} \approx 2T_{C,0}$ ). These relations are summarized in Table 8. This indicates that at least some dust in the original shell survived the outburst and has heated up.

The UVOT data trace the evolution of the hotter blackbody detected in the initial outburst. The blackbody radius decreases from about 1.7 to 0.55 au over the 10 days of observations while remaining at a steady temperature of  $\approx 9200\ \text{K}$ , as shown in Figure 5. This is consistent with a receding photosphere of the initial outburst. We can use this observation to constrain the radius of the progenitor/surviving star to  $\lesssim 0.55\ \text{au}$ , since we expect the photosphere of the eruption to exceed the radius of the star at all times.

### 3.1.3. The Progeny

About 200 days after its discovery, SN 2010da dips to  $m_i \approx 23\ \text{mag}$  in the optical but returns to  $m_i \approx 20\ \text{mag}$  after 500 days. In the same time frame, the IR flux declines by about one magnitude to  $m_{3,6} \approx 18.7\ \text{mag}$  at 460 days. The optical light curve then appears to settle into an aperiodic, variable state that oscillates between  $m_{r,i} \approx 20\ \text{mag}$  and  $m_i \approx 22\ \text{mag}$  every 500–1000 days. The IR light curve remains roughly at its pre-eruption brightness, but the color becomes much bluer (from  $m_{3,6} - m_{4,5} \approx 0.2$  to  $\approx -0.2$ ). Beginning around  $\approx 1500$  days after discovery, the IR light curve begins to rise to magnitudes comparable to the initial outburst. We refer to the surviving star as the *progeny* of SN 2010da.

The progeny’s optical/IR SED can be roughly described by a single blackbody with variable excess flux in the optical. After 500 days, the NIR and IR fluxes are fit by a blackbody with a radius of  $\approx 6\ \text{au}$  and a temperature of  $\approx 2000\ \text{K}$ . The derived radius is smaller than the progenitor radius at  $\approx 10\ \text{au}$ , and the temperature is higher than the blackbody temperature of the pre-eruption SED ( $T = 1500\ \text{K}$ ), consistent with the color

shift seen in the IR. The optical flux, however, varies by  $\approx 2\ \text{mag}$  even two years after the initial outburst. Fitting our NIR/IR measurements to blackbodies, we track the bolometric luminosity of the system over time, as shown in Figure 6. The luminosity of the progeny and its surrounding environment is about two to five times larger than the progenitor of SN 2010da excluding contribution from the UV/optical, which supplies  $\sim 10\%\text{--}20\%$  of the total luminosity.

We compare the SEDs taken more than 800 days after the initial outburst to a variety of SEDs of massive stars in the LMC analyzed by Bonanos et al. (2009). We group these massive stars by their spectroscopic classification reported by Bonanos et al. (2009), and we construct “typical” SED ranges for each filter within each group. The SEDs for red, yellow, and blue supergiants (RSGs, YSGs, BSGs), LBVs and B[e] stars compared to the SED of the progeny of SN 2010da are shown in Figure 7. Here, we are defining B[e] stars as any star with B[e] emission lines (e.g., Hydrogen Balmer, iron, etc.), regardless of luminosity class. The progeny’s SED most closely matches the SED of a typical RSG. As a test, we also convert our SDSS bandpasses into Johnson magnitudes and search for the nearest neighbor of the progeny SED within the space of the magnitudes used by Bonanos et al. (2009); the nearest neighbor is [SP77]46-31-999, an M2 Iab star. The fact that the SED of the progeny most closely resembles that of an RSG does not necessarily imply that the progenitor or progeny is an RSG. In fact, the small radius we infer from the *Swift*/UVOT data ( $\lesssim 0.5\ \text{au}$ ) rules out most RSG candidates. Both broadband photometry and spectroscopy are necessary when classifying obscured, massive stars.

### 3.2. Spectroscopic Properties of SN 2010da

Throughout our observations, spanning from 2 to 1881 days after the optical discovery, the spectra of SN 2010da exhibit strong hydrogen Balmer and Paschen, He I and II, Fe II, and Ca II emission lines. Early spectra reveal P-Cygni profiles in the Balmer, Paschen, and helium lines, while later spectra develop strong nebular emission lines. A full list of these lines with a  $3\sigma$  detection in at least one epoch and their properties is provided in Table 9. The low-resolution spectra are shown in Figure 8, and the high-resolution spectra are shown in Figure 9. The high-resolution spectra have been normalized by fitting a low-order polynomial to the smoothed spectra. The strong Balmer lines, low excitation emission lines (especially Fe II), the forbidden lines, and the IR excess all indicate that the progeny of SN 2010da exhibits B[e] phenomena by the criteria enumerated in Lamers et al. (1998). This classification scheme is purely observational but is linked to a complex CSM surrounding the star (see Lamers et al. 1998 Section 2.2). We observe the development of high-ionization emission lines of iron at later epochs and continuous He II 4686 Å emission. Both of these observations are due to the presence of a hard

**Table 5**  
Ground-based Optical Spectroscopy

Date (UT)	Epoch (days)	Instrument	Exposure (s)	Grating/Grism	Resolution (Å)
2010 May 25	2	GMOS-S	1200	R400	4
2010 May 25	2	GMOS-S	600	B600	4
2010 Jun 07	15	MIKE	1800	R2	0.3
2010 Jun 07	15	MIKE	1800	R2.4	0.4
2010 Jul 02	40	GMOS-S	1200	R400	4
2010 Jul 02	40	GMOS-S	900	B600	4
2011 Oct 21	516	MagE	1200	...	2
2011 Dec 29	585	IMACS	300	300-17.5	5
2015 May 16	1819	IMACS	1500	300-17.5	5
2015 Jul 17	1881	IMACS	1800	300-17.5	5

radiation field (UV/X-ray emission) associated with the HMXB nature of the object (see Section 4.3).

### 3.2.1. Hydrogen Balmer Lines

The Balmer lines exhibit some of the most drastic changes of the spectrum over the span of our observations. Their equivalent widths roughly follow the optical/IR flux variations and appear to be significantly increasing in the most recent observations. A time sequence of the Balmer lines is shown in Figure 10.

The spectra are marked by large  $H\alpha$  luminosity, contributing up to  $\approx 30\%$  of the  $r$ -band flux at later times. Following the initial outburst,  $H\alpha$  is well described by a Lorentzian profile with a full width at half maximum of  $\approx 560 \text{ km s}^{-1}$ . The full width at continuum intensity is  $\approx 3600 \text{ km s}^{-1}$ . The  $H\alpha$  flux immediately following the discovery is  $\approx 6.6 \times 10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$  and approximately halves 40 days later. As the object cools, the continuum flux decreases while the  $H\alpha$  flux remains relatively constant at  $\approx (2-3) \times 10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2}$ . At the same time,  $F_{H\alpha}/F_{H\beta}$  increases from  $\approx 4$  to  $\approx 8$  in the first 40 days. This is far greater than the expected value of  $F_{H\alpha}/F_{H\beta} \approx 2.8$  for Case B recombination at  $\sim 10^4 \text{ K}$ , the approximate temperature of the hotter blackbody component in the SED during outburst. While dust extinction may account for this excess, the continuum is unabsorbed. An alternative possibility is that a high-density CSM affects the Balmer decrement via self-absorption and collisions (Drake & Ulrich 1980). Using the static slab model at  $10^4 \text{ K}$  from Drake & Ulrich (1980), we find that the observed  $F_{H\alpha}/F_{H\beta}$  ratio roughly corresponds to a density of  $n_e \sim 10^{10}-10^{12} \text{ cm}^{-3}$ . At these densities,  $F_{H\alpha}/F_{H\beta}$  is suppressed to  $\approx 0.3$ ,  $F_{H\gamma}/F_{H\beta}$  to  $\approx 0.2$ , and  $F_{H\epsilon}/F_{H\beta}$  to  $\approx 0.15$ . The observed line fluxes roughly match these predictions during the initial outburst. 40 days later, the  $F_{H\alpha}/F_{H\beta}$  ratio remains consistent with  $n_e \sim 10^{10}-10^{12} \text{ cm}^{-3}$ .

In our high-resolution MIKE spectrum taken 14 days after discovery, the  $H\alpha$  line includes multiple components. We fit the  $H\alpha$  profile with three Gaussian components: a narrow component (FWHM  $\approx 70 \text{ km s}^{-1}$ ), an intermediate component (FWHM  $\approx 500 \text{ km s}^{-1}$ ), and a broad component (FWHM  $\approx 1060 \text{ km s}^{-1}$ ) with  $\chi_r^2 \approx 1.9$  (see Figure 11). Multi-component (specifically three-component) lines are common in dusty ILOTs (Turatto et al. 1993; Berger et al. 2009; Van Dyk & Matheson 2012; Tartaglia et al. 2015). The narrow component is broader than the other detected narrow emission lines (e.g., other Balmer lines with FWHM  $\sim 40-60 \text{ km s}^{-1}$ ) possibly due to electron scattering.

These narrow components are consistent with a pre-existing wind, possibly from an earlier red supergiant phase. As with SNe IIn emission lines, the intermediate component is ascribed to the shockwave-CSM interaction (see Chevalier & Fransson 1994), though the velocity is nearly an order of magnitude slower than in SNe. The intermediate component is significantly redshifted (by  $\approx 140 \text{ km s}^{-1}$ ) relative to the narrow component. The apparent redshift may be an artifact of electron scattering through high optical depths and is often seen in other dusty ILOTs, giants, Wolf-Rayet stars, and other stars experiencing significant mass loss (Hillier 1991; Humphreys et al. 2011). The broadest component is only identified in  $H\alpha$ , which may be due to lower signal-to-noise in the other lines or additional scattering. The central wavelength of this component is consistent with that of the intermediate component, also suggesting a common physical origin.

At later times, the  $H\alpha$  emission line can be described as roughly Gaussian or Lorentzian with an extended red wing and an FWHM  $\sim 300-600 \text{ km s}^{-1}$ . This red wing is especially apparent in the high-resolution MagE spectrum at 516 days. The FWHM of the late-time  $H\alpha$  emission is consistent with the intermediate component of the  $H\alpha$  line during the initial eruption, while the narrow component remains unresolved in all other spectra. This suggests that the late-time emission is powered by a persistent wind or mass loss consistent with that of a blue or yellow supergiant.

### 3.2.2. Ca II Lines

Narrow [Ca II] lines (FWHM  $\lesssim 50 \text{ km s}^{-1}$ ) are detected in our highest resolution (MIKE) spectrum 15 days after discovery and possibly again in the MagE spectrum (Figure 12). Similar forbidden calcium emission was observed in NGC 300 OT2008-1 (Berger et al. 2009) and SN 2008S (Botticella et al. 2009), as well as in several warm hypergiants (Humphreys et al. 2013), and its presence is typically associated with a dusty environment. Because collisional de-excitation normally drives calcium to its ground state, the [Ca II] doublet is associated with cooler, low-density regions. Forbidden calcium lines are additionally suppressed by UV radiation due to the low ionization potential of calcium. Highly ionized iron and He II lines indicate a strong UV radiation field. We can conclude from this fact and the narrow line shape that the [Ca II] forbidden lines are from excited calcium located in the outer CSM, likely in the original dust shell at  $\approx 10 \text{ au}$ .

The presence of the calcium triplet also supports the existence of a cool, low-density circumstellar environment (Polidan & Peters 1976). However, the calcium triplet is significantly blended with hydrogen Paschen emission, so we

**Table 6**  
*Swift UVOT/Photometry*

Date (UT)	Epoch	AB Magnitude					
		UVW2	UVM2	UVW1	U	B	V
2010 May 26	3	18.27 ± 0.08	18.06 ± 0.07	17.48 ± 0.07	16.72 ± 0.07	16.48 ± 0.07	16.23 ± 0.07
2010 May 27	4	18.20 ± 0.08	17.96 ± 0.07	17.43 ± 0.071	16.67 ± 0.07	16.43 ± 0.07	16.17 ± 0.08
2010 May 28	5	18.33 ± 0.08	18.10 ± 0.07	17.57 ± 0.07	16.73 ± 0.067	16.53 ± 0.07	16.31 ± 0.08
2010 May 31	8	18.89 ± 0.09	18.58 ± 0.08	17.97 ± 0.076	17.21 ± 0.07	16.92 ± 0.08	16.61 ± 0.09
2010 Jun 01	9	19.31 ± 0.10	18.79 ± 0.08	18.25 ± 0.08	17.46 ± 0.08	17.20 ± 0.08	17.00 ± 0.11
2010 Jun 03	11	19.43 ± 0.11	18.93 ± 0.09	18.41 ± 0.10	17.48 ± 0.08	17.45 ± 0.10	17.05 ± 0.11
2010 Jun 05	14	19.61 ± 0.11	19.30 ± 0.09	18.61 ± 0.085	17.77 ± 0.08	17.63 ± 0.10	17.35 ± 0.12
2010 Jun 07	16	19.89 ± 0.12	19.58 ± 0.10	18.81 ± 0.092	18.06 ± 0.09	17.93 ± 0.11	17.75 ± 0.16
2010 Jun 09	18	20.07 ± 0.13	19.62 ± 0.10	1.46 ± 0.0	18.24 ± 0.10	18.14 ± 0.13	17.78 ± 0.16
2010 Jun 11	20	20.13 ± 0.14	19.83 ± 0.11	19.11 ± 0.10	18.53 ± 0.11	18.56 ± 0.17	18.02 ± 0.20
2010 Jun 16	24	20.49 ± 0.16	20.18 ± 0.13	19.46 ± 0.123	18.71 ± 0.12	18.87 ± 0.21	18.48 ± 0.27
2010 Jun 17	25	20.53 ± 0.17	20.10 ± 0.15	19.61 ± 0.15	18.92 ± 0.14	18.85 ± 0.20	18.59 ± 0.32
2010 Jun 18	26	20.51 ± 0.16	20.15 ± 0.13	19.52 ± 0.12	19.08 ± 0.14	19.03 ± 0.23	>19.01
2010 Jun 21	29	21.05 ± 0.26	20.57 ± 0.24	19.89 ± 0.148	19.02 ± 0.14	>19.71	>18.25
2010 Jun 26	34	20.8 ± 0.24	20.08 ± 0.22	19.84 ± 0.181	19.28 ± 0.20	>19.42	>17.98
2010 Jul 02	40	21.21 ± 0.25	21.13 ± 0.23	20.54 ± 0.219	19.76 ± 0.21	>19.82	>19.00
2011 Oct 26	521	>21.46	>20.86	>20.72	20.84 ± 0.30	>19.43	>18.62
2011 Oct 27	522	>22.32	>21.02	>20.84	>20.30	>19.53	>18.72
2011 Oct 28	523	>21.64	21.85 ± 0.33	>20.92	>20.39	>19.60	>18.80

cannot make a definitive statement about the line shape or strength. The Ca H and K doublet, typically associated with the calcium triplet, are also blended with H $\eta$  and an iron line.

### 3.2.3. Fe Lines

We detect strong Fe II emission lines in all spectra. The Fe II features roughly match the hydrogen Balmer series in shape, FWHM, and line offset, indicating that these features also arise from material within the ejecta and CSM. Fe II emission is seen in NGC300 OT2008-1 and M85 OT2006-1 (Berger et al. 2009; Bond et al. 2009), though the lines seen in SN 2010da are notably stronger.

In addition to Fe II, we detect emission from high ionization, forbidden iron lines, including [Fe VII], [Fe X] and [Fe XI] in the last two epochs of spectroscopy (Epochs 1819 and 1881; see Figure 13). These forbidden iron lines are not typically seen in ILOTs due to their weakly ionizing radiation. High-ionization iron lines are occasionally found in SNe II $n$  such as SN 1997eg (Hoffman et al. 2008) and SN 2005ip (Smith et al. 2009) due to shock heating of the surrounding CSM. Unlike SNe II $n$ , the iron lines seen in SN 2010da do *not* arise from continual shock heating over hundreds of days. Instead, these lines arise in regions of diffuse gas surrounding the progeny, which are heated to temperatures of about  $2 \times 10^6$  K, the approximate ionization temperature for lines such as [Fe XII] and [Fe XIV] (Corliss & Sugar 1982), by X-rays from the compact companion (see Section 4.3).

### 3.2.4. He Lines

He I and II emission lines are seen throughout our observations, as shown in Figure 14. The widths (FWHM  $\sim 200$ – $400$  km s $^{-1}$ ) and profiles of the He I lines largely follow the Balmer series with a double-peaked structure in our high-resolution spectra. We additionally detect single-peaked He II 4686 Å emission during each epoch. The low-resolution spectra are unable to resolve the He II 4686 Å line, but our MIKE spectrum reveals an FWHM  $\approx 270$  km s $^{-1}$ . He II 4686 Å has a relatively high-ionization potential and is

sensitive to the EUV flux of the system. For this reason, it is often linked to the accretion onto a compact object (Lewin et al. 1997). The continual presence of He II 4686 Å emission in each of our observations is due to the compact companion and indicates consistent mass transfer onto this compact object.

### 3.2.5. Additional Absorption and Emission Features

Hydrogen Paschen emission lines extending to approximately Pa30 are observed within the first 15 days of the outburst. The high-resolution MIKE spectrum reveals double-peaked emission with narrow and broad components, similar to the Balmer lines. These profiles are not resolved in our lower-resolution spectra. Within the first 40 days, the ratio  $F(P_\delta)/(H_\beta)$  decreases from  $\approx 0.2$  to  $\approx 0.08$ . The latter value is roughly consistent with that expected from Case B recombination ( $\approx 0.07$ ), though the effect of high electron density on this line is unclear.

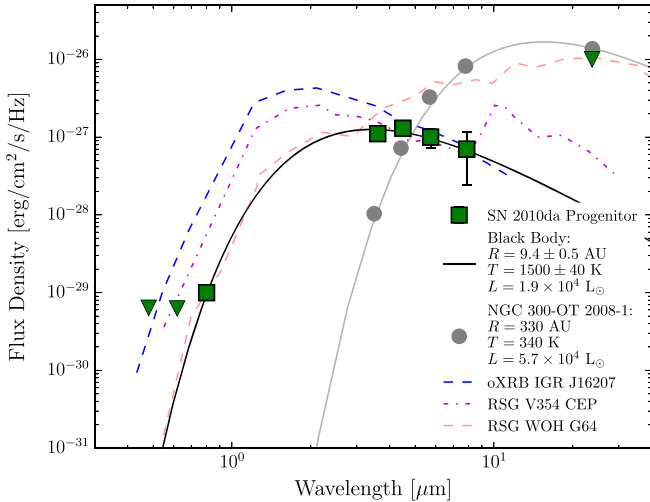
At early times, we detect O I at 7774 Å and 8446 Å in emission. The O I 8446 Å line shows the same double-peaked profile as the Balmer series, while O I 7774 Å maintains a P-Cygni profile until 40 days after the initial eruption. After 1800 days, O I 7774 Å becomes undetectable while O I 8446 Å strengthens. The expected ratio between these lines is  $F(\text{O I } 8446 \text{ \AA})/F(\text{O I } 7774 \text{ \AA}) = 0.6$ , implying that O I 7774 Å should be detectable. The independent strengthening of O I 8446 Å can be attributed to Ly $\beta$  emission, which is outside of our observed spectral range. Ly $\beta$  photons pump O I from the ground state to an unstable state whose decay produces O I 8446 Å emission (Mathew et al. 2012). This is consistent with the increased Balmer emission and UV flux at later times. In addition to O I, we detect [O I] and [O III] features. Unfortunately, the [O III] 4363 Å feature appears to be blended with either [Fe IX] or an Fe II emission line, making it difficult to use the [O III] ratios to calculate the electron temperature.

Unresolved Na I D lines are observed in the two latest epochs (517 and 1817 days) as emission and absorption respectively. The variability of these lines indicates that they are associated with the CSM rather than the interstellar medium.

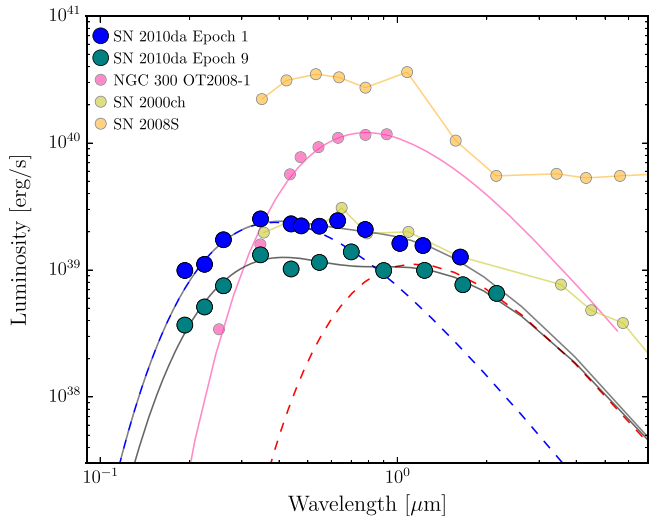


**Table 7**  
Chandra Photometry

Start Date (UT)	Epoch	Proposal ID	PI	Counts	Detection Significance ( $10^{-15}$ erg s $^{-1}$ cm $^{-2}$ )	0.3–10 keV Flux
2010 Sep 24	123	12238	Williams	77 ± 9	21σ	50.7 $^{+8.9}_{-9.2}$
2014 May 16	1453	16028	Binder	7 ± 3	3σ	5.6 $^{+2.4}_{-3.3}$
2014 Nov 17	1638	16029	Binder	140 ± 12	52σ	142 $^{+12}_{-28}$



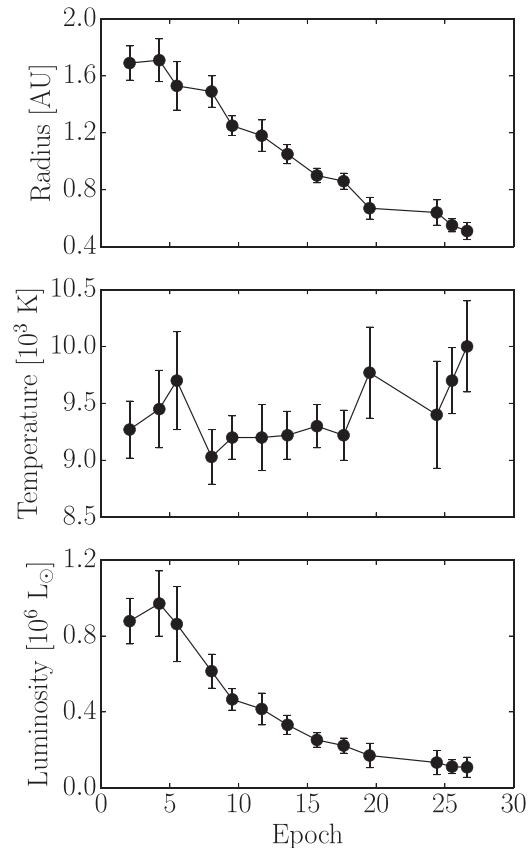
**Figure 3.** Spectral energy distribution of the progenitor of SN 2010da (green squares) with a blackbody fit (black). Also shown are an obscured HMXB (blue dashed line, IGR J16207-5129; Tomsick et al. 2006) and two RSGs (purple and pink lines, see Section 3.1.1 for a discussion of these objects; van Loon et al. 1999; Mauron & Josselin 2011). The latter three objects have been arbitrarily scaled to match the luminosity of the progenitor of SN 2010da. NGC 300’s other well-known impostor, NGC 300 OT2008-1, as well as its best-fit blackbody, are shown in gray for comparison. Although NGC 300 OT2008-1 and SN 2010da are spectroscopically similar, the progenitor of SN 2010da is obscured by significantly warmer dust. Downward facing triangles are 3σ upper limits.



**Figure 4.** Spectral energy distribution of SN 2010da one day (dark blue points) and nine days (dark cyan points) after discovery. Both SEDs are fit with a two-component blackbody model. The total fit is shown in black, while the components of the outburst SED model are shown in blue and red for the first epoch. Also shown are the SEDs of three similar ILOTs during outburst, NGC 300 OT2008-1 (pink), SN 2000ch (yellow), and SN 2008S (orange). NGC 300 OT2008-1 and SN 2008S are ILOTs with very red progenitors, while SN 2000ch is an LBV-like star.

**Table 8**  
Summary of SN 2010da Blackbody Fits

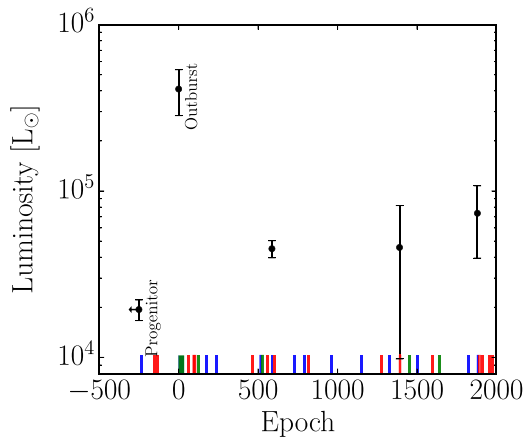
	Progenitor	Epoch 1	Epoch 9
$T_C$	1500 ± 40 K	3230 ± 490 K	2760 ± 250 K
$T_H$	...	9440 ± 280 K	9080 ± 330 K
$R_C$	9.4 ± 0.5 au	9.5 ± 2.9 au	10.5 ± 1.6 au
$R_H$	...	1.59 ± 0.14 au	1.25 ± 0.13 au



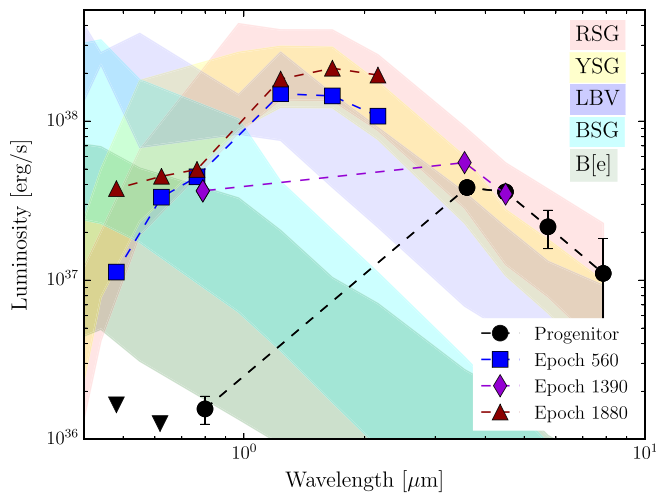
**Figure 5.** We fit the *UVOT* data set to a blackbody as a function of time. We find that the blackbody radius recedes (top), while the temperature remains roughly constant (middle). The estimated luminosity decreases with the receding radius (bottom).

### 3.3. X-Ray Spectral Modeling

We model the X-ray emission from SN 2010da and its progeny using XSPEC version 12.8.2n (Arnaud 1996). We use the Cash statistic, a derivative of the Poisson likelihood, as our fit statistic. To test our fits, we use the XSPEC built-in command `goodness` to perform  $10^4$  Monte Carlo simulations of the spectral data. For each simulation, the program calculates the Cramér von Mises (CVM) test statistic, which is shown to be a good fit statistic for the data derived from a Poisson



**Figure 6.** Bolometric luminosity of SN 2010da as derived from the NIR/IR flux as a function of time (black). The colored lines at the bottom of the plot indicate our photometric X-ray (green), optical (blue), and infrared (red) coverage for reference. The progeny of SN 2010da system has a consistently higher bolometric luminosity than its progenitor.



**Figure 7.** Spectral energy distributions of SN 2010da compared to SEDs calculated using massive stars in the LMC (Bonanos et al. 2010). The red region is the typical range for an RSG (M-type), yellow for a YSG (K-type), cyan for a BSG (B-type), dark blue for an LBV, and dark green for a B[e] star. Although SN 2010da displays B[e] properties in its spectrum, its SED seems to follow that of a RSG or YSG. Downward facing triangles are  $3\sigma$  upper limits.

distribution (Spinelli & Stephens 1997). If about 50% of these simulations have a CVM statistic less than that of our model, the best-fit model is a good representation of the data. A percentage much lower than 50% implies that our model is over parametrized, and a percentage much greater than 50% implies that our model is inconsistent with the data. All reported errors correspond to  $1\sigma$  error bars (the 68% confidence interval).

We combine all of the *Swift*/XRT 0.5–8 keV data taken within 40 days of the outburst and fit it to a power law with Galactic absorption (tbabs \* pow) with  $N_H$ ,  $M_W = 4 \times 10^{20} \text{ cm}^{-2}$ . We find that additional absorption over-parametrizes the model (goodness = 15%), but an excess column density as large as  $N_H \approx 5 \times 10^{21} \text{ cm}^{-2}$  is consistent with the data. Our best-fit model is described by  $\Gamma = -0.05^{+0.11}_{-0.10}$  with an unabsorbed 0.3–10.0 keV flux of  $9.62^{+0.87}_{-0.85} \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2}$  (assuming  $N_H = 0$ ). This corresponds to a luminosity of  $3.98^{+0.36}_{-0.35} \times 10^{38} \text{ erg s}^{-1}$ .

Similarly, we fit the first *Chandra* observation (Epoch 123) to an absorbed power law. We find the best-fit model is described by  $\Gamma = 0.26^{+0.20}_{-0.21}$  with an unabsorbed 0.3–8.0 keV luminosity of  $1.95^{+0.17}_{-0.48} \times 10^{37} \text{ erg s}^{-1}$ . We find an absorption upper limit beyond the Galactic column of  $N_H \lesssim 4 \times 10^{21} \text{ cm}^{-2}$ . The estimated column density and the photon index are degenerate such that a higher column density implies a softer power law.

Due to limited statistics, we are unable to fit a spectrum to the second *Chandra* observation (Epoch 1453). In the third *Chandra* observation (Epoch 1638), there is a significant decrease in counts between  $\approx 2$ –3 keV. We are unable to fit this spectrum to a single power-law or blackbody component and instead combine a power law with either a soft Bremsstrahlung (bremss) or blackbody disk (diskbb) model, with no statistical preference for either model based on the CVM statistic. For both models, we obtain a similar power law with an index of  $\Gamma = -2.2^{+0.3}_{-0.5}$  for the Bremsstrahlung model and  $\Gamma = -2.3^{+0.4}_{-0.4}$  for the disk model. The Bremsstrahlung component has a temperature of  $0.6^{+0.3}_{-0.2} \text{ keV}$ , while the disk model has an inner-disk temperature of  $0.33^{+0.08}_{-0.06} \text{ keV}$ . While these fits were performed by fixing the hydrogen column density to the Galactic value, fixing  $N_H$  to values as high as  $4 \times 10^{21} \text{ cm}^{-2}$  also gives reasonable (though statistically less favorable) fits with softer power laws. This hard power law differs from the recent results of Binder et al. (2016), who find  $\Gamma \approx 0$ . Specifically, we are unable to reconcile the double peak in the spectrum with softer power laws (see Figure 15). The extremely hard power law in our models indicates that additional and detailed modeling is necessary to explain this unusual *Chandra* spectrum. All X-ray spectra and models are shown in Figure 16.

### 3.4. The X-Ray and UV Light Curves

Using archival observations of SN 2010da (Section 2.7), we are able to construct the X-ray and UV light curves of SN 2010da. The full X-ray light curve is shown in Figure 17. We build the *Swift*/XRT light curve by converting the light curve produced automatically by the UK *Swift* Science Data Center from a count rate to a flux using the conversion factor found for the XRT spectrum. This light curve was dynamically binned using a rate factor of 10 and a bin factor of 5.

Binder et al. (2011) estimate a  $3\sigma$  upper limit on the unabsorbed 0.3–10 keV luminosity of the progenitor to be  $\approx 3 \times 10^{36} \text{ erg s}^{-1}$  using archival *XMM-Newton* data. However, about 1300 days before the optical outburst, we find a weak *Swift*/XRT detection at  $2.6\sigma$  with a luminosity of  $1.8^{+0.7}_{-0.7} \times 10^{37} \text{ erg s}^{-1}$ , indicating X-ray variability even before the 2010 optical outburst.

During the transient, the X-ray luminosity increases to a peak of  $\approx 6 \times 10^{39} \text{ erg s}^{-1}$ , making SN 2010da an ultra-luminous X-ray source well above the Eddington limit of a  $1.4 M_\odot$  neutron star. (We note that this luminosity is larger than the luminosity reported from the spectral fit in Section 3.3 and Binder et al. 2011, because the spectral fit averaged the luminosity over 40 days following the initial outburst.) In the week following discovery, the X-ray luminosity fluctuates between  $2 \times 10^{39} \text{ erg s}^{-1}$  and  $6 \times 10^{39} \text{ erg s}^{-1}$  before decaying with an  $e$ -folding time of  $\approx 3.5$  days. This decay rate is slightly longer than that of the UVOT light curves ( $\approx 10$  days, shown in Figure 18) and is much shorter than the decay rates

**Table 9**  
Catalog of Spectral Lines Identified in Spectra of SN 2010da

Line	Epoch	Profile	Line Center (Å)	EW (Å)	FWHM (km s <sup>-1</sup> )	Flux (10 <sup>-16</sup> erg s <sup>-1</sup> cm <sup>-2</sup> )
Fe I 3752.88	2	SP	3752.4	-0.2 (0.1)	≲400	7 (1)
Fe I 3774.76	2	SP	3772.1	-1.30 (0.02)	420 (40)	6.46 (0.08)
Fe I 3800.61	2	DP	3802.0, 3804.0	-2.21 (0.09)	<400 450	10.56 (0.09)
	15	SP	3800.1	-4 (1)	420 (10)	...
	40	SP	3800.0	-1.4 (0.3)	≲300	0.5 (0.3)
H $\eta$	2	SP	3839.0	-1.83 (0.08)	570 (60)	7 (2)
	15	SP	3838.2	-1.8 (0.8)	440 (20)	...
	40	SP	3836.9	-2.4 (0.2)	140 (50)	1.0 (0.4)
H $\zeta$	2	P-Cyg	3885.3, 3891.9	...	...	...
	15	DP	3890.1 3891.8	-9 (3)	430 (10) 65 (9)	...
	40	SP	3891.2	-3.1 (0.2)	350 (40)	2.6 (0.2)
Fe I 3935.81 and Ca II K	2	P-Cyg	3932.9, 3937.9	...	...	...
	15	SP	3935.77	-5 (1)	390 (10)	...
	40	SP	3936.97	-15 (2)	470 (30)	1.1 (0.3)
H $\epsilon$ and Ca II H	2	DP	3976.0 3973.13	-3.8 (0.1)	1100 (300) 490 (40)	27 (1)
	15	DP	3971.8 3972.8	-10 (2)	490 (20) 130 (20)	...
	40	SP	3971.36	-18 (1)	<300	5.2 (0.3)
H $\delta$	2	SP	4104.72	-5.52 (0.07)	560 (20)	32 (1)
	15	DP	4102.9 4104.1	-6 (1)	42 (2) 449 (8)	...
	40	SP	4103.76	-12 (1)	480 (30)	2.2 (0.5)
	516	SP	4014.11	-12 (4)	110 (10)	...
Fe II 4174.62	2	ND	...	...	...	...
	15	SP	4174.86	-0.23 (0.05)	100 (10)	...
	40	SP	4174.81	-5.2 (0.2)	220 (30)	2.5 (0.3)
	516	SP	4175.29	-5.0 (0.7)	150 (20)	...
Fe II 4180.03	2	ND	...	...	...	...
	15	SP	4180.15	-0.7 (0.1)	220 (30)	...
	40	SP	4179.66	-15 (3)	590 (70)	1.3 (0.3)
	516	SP	4180.15	-6 (2)	900 (200)	...
Fe I 4237.12	2	SP	4236.24	-0.75 (0.02)	500 (50)	4.3 (0.7)
	15	SP	4234.9	-2.2 (0.2)	109 (6)	...
	40	SP	4234.6	-7.8 (0.6)	350 (50)	1.6 (0.2)
	516	SP	4235.0	-8 (1)	450 (60)	...
H $\gamma$	2	SP	4342.9	-8.8 (0.1)	590 (20)	53 (1)
	15	DP	4342.2 4342.8	-11 (2)	458 (9) 56 (2)	...
	40	SP	4342.0	-29 (4)	180 (10)	8.6 (0.2)
	516	DP	4341.4 4342.9	-19 (3)	510 (200) 120 (10)	...
	1819	SP	4343.2	-9 (3)	400 (60)	2.2 (0.3)
	1881	SP	4343.2	-20 (20)	630 (140)	7 (1)
[O III] 4363 and [Fe IX] 4359 and Fe II	2	ND	...	...	...	...
	15	SP	4364.3	-0.027 (0.008)	40 (4)	...
	40	ND	...	...	...	...
	516	ND	...	...	...	...
	1819	SP	4365.8	-5 (1)	260 (50)	2.2 (0.1)
	1881	SP	4365.9	-30 (20)	500 (100)	7.2 (0.2)
Fe I 4384.77	2	P-Cyg	4382.0 4388.3	...	...	...
	15	SP	4387.1	-0.8 (0.1)	280 (20)	...
	40	SP	4387.0	-2.8 (0.2)	460 (80)	1.4 (0.1)
	516	SP	4388.0	-3.3 (0.4)	360 (40)	...
	1819	ND	...	...	...	...
Fe I 4416.36	2	SP	4418.6	-0.27 (0.01)	340 (80)	3.1 (0.5)
	15	SP	4418.5	-2.0 (0.2)	280 (10)	...
	40	SP	4417.3	-3.1 (0.3)	450 (70)	0.8 (0.3)

**Table 9**  
(Continued)

Line	Epoch	Profile	Line Center (Å)	EW (Å)	FWHM (km s <sup>-1</sup> )	Flux (10 <sup>-16</sup> erg s <sup>-1</sup> cm <sup>-2</sup> )
He I 4471.5	516	SP	4418.0	-3.8 (0.4)	390 (40)	...
	1819	ND	...	...	...	...
	1881	ND	...	...	...	...
	2	P-Cyg	4467.0	...	...	...
	2		4474.5			
	15	SP	4473.4	-2.7 (0.3)	370 (10)	...
	40	SP	4473.3	-3.7 (0.3)	340 (50)	0.8 (0.1)
Fe I 4585	516	SP	4474.6	-2.9 (0.6)	250 (40)	...
	1819	ND	...	...	...	...
	1881	ND	...	...	...	...
	2	SP	4586.0	-1.30 (0.03)	340 (80)	6 (1)
	15	SP	4586.0	-7 (1)	330 (50)	...
	40	SP	4585.9	-2.4 (0.7)	440 (50)	1.6 (0.7)
	516	SP	4585.9	-1.4 (0.3)	130 (20)	...
He II 4686	1819	SP	4584.8	>-6	≤230	<0.9
	1881	SP	4587.8	>-6	600 (100)	<1.0
	2	SP	4683.0	-3.2 (0.6)	550 (30)	23 (1)
	15	SP	4687.2	-3.0 (0.2)	270 (10)	...
	40	SP	4686.7	-7.5 (0.5)	330 (20)	2.7 (0.2)
	516	SP	4687.6	-7 (2)	270 (20)	...
	1819	SP	4687.5	-12 (4)	480 (60)	1.6 (0.3)
Hβ	1881	SP	4687.6	-30 (10)	560 (70)	4.5 (0.7)
	2	SP	4863.6	-20.9 (0.4)	530 (20)	160.8 (0.7)
	15	DP	4862.9	-16 (1)	66 (3)	...
			4864.6		499 (7)	
	40	SP	4863.4	-110 (10)	360 (10)	36.3 (0.2)
	516	DP	4863.7	-70 (10)	106 (4)	...
			4861.7		540 (60)	...
[O III] 4959	1819	SP	4863.8	-80 (20)	450 (10)	12.5 (0.2)
	1881	SP	4864.0	-170 (4)	620 (10)	24.7 (0.3)
	2	ND	...	...	...	...
	15	SP	4960.3	>-0.4	50 (10)	...
	40	SP	4960.0	-2.3 (0.1)	280 (30)	0.4 (0.1)
	516	ND	...	...	...	...
	1819	SP	4960.9	-8 (1)	300 (40)	1.8 (0.3)
[O III] 5007	1881	SP	4958.0	-20 (7)	530 (80)	3.1 (0.3)
	2	SP	5007.9	-0.22 (0.03)	<240	2.1 (0.5)
	15	SP	5008.0	-1.4 (0.1)	63 (2)	...
	40	SP	5008.0	-4.4 (0.5)	≤240	1.7 (0.3)
	516	ND	...	...	...	...
	1819	SP	5009.1	-25 (7)	350 (20)	5.6 (0.3)
	1881	DP	5008.3	-80 (10)	≤300	9.4 (0.3)
Fe I 5017.87			5012.3		500 (20)	
	2	SP	5020.15	-2.36 (0.02)	440 (20)	17.6 (0.2)
	15	DP	5021.1	-5.1 (0.8)	430 (20)	2.9 (0.2)
			5020.0		37 (3)	
	40	SP	5019.8	-7 (1)	350 (50)	3.2 (0.1)
	516	DP	5019.7	-7 (1)	420 (50)	...
			5020.7		≤90	
[Fe VII] 5159 and Fe II	1819	ND	...	...	...	...
	1881	ND	...	...	...	...
	2	SP	5159.8	>0.07	170 (50)	1.1 (0.1)
	15	SP	5160.2	-1.2 (0.4)	<80	...
	40	ND	...	...	...	...
	516	SP	5160.6	<0.6	130 (30)	...
	1819	SP	5161.0	-6 (1)	<290	0.4 (0.1)
[Fe VII] 5276	1881	SP	5161.0	-5 (2)	400 (100)	1.0 (0.2)
	2	-	...	...	...	Note: on chip
	15	SP	5277.4	-0.4 (0.1)	300 (30)	...
	40	-	...	...	...	Note: on chip
	516	SP	5276.6	-2.9 (0.4)	60 (10)	...
	1819	SP	5277.7	>-4	<400	<0.7
	1881	SP	5274.0	>-4	<600	<2

**Table 9**  
(Continued)

Line	Epoch	Profile	Line Center (Å)	EW (Å)	FWHM (km s <sup>-1</sup> )	Flux (10 <sup>-16</sup> erg s <sup>-1</sup> cm <sup>-2</sup> )
[O I] 5577	2	ND	...	...	...	...
	15	ND	...	...	...	...
	40	ND	...	...	...	...
	516	SP	5576.5	-0.08 (0.04)	60 (20)	...
	1819	SP	5577.2	-11 (2)	180 (20)	0.20 (0.08)
	1881	SP	5577.3	+4 (1)	300 (50)	-1.0 (0.1)
[N II] 5755 and Fe II	2	ND	...	...	...	...
	15	SP	5756.3	-1.2 (0.1)	200 (100)	0.4 (0.1)
	40	ND	...	...	...	...
	516	ND	...	...	...	...
	1819	SP	5756.5	-0.9 (0.5)	200 (100)	0.29 (0.03)
	1881	SP	5753.6	-6 (2)	800 (200)	1.5 (0.4)
He I 5877	2	SP	5890.0	-5.1 (0.1)	450 (20)	36.6 (0.7)
	15	DP (LG)	5877.3	-6.2 (0.8)	30 (4)	...
			5878.7	...	370 (10)	...
	40	SP	5878.2	-7.6 (0.7)	290 (20)	6.7 (0.3)
	516	SP	5877.5	-8 (2)	170 (20)	...
	1819	SP	5878.0	-11 (9)	360 (80)	0.4 (0.3)
[Fe VII] 6086	1881	SP	5878.5	-32 (7)	500 (40)	4.3 (0.5)
	2	SP	6088.4	-0.07 (0.01)	260 (60)	0.2 (0.1)
	15	SP	6088.3	>-0.3	240 (30)	...
	40	SP	6087.6	0.83 (0.03)	230 (30)	0.60 (0.06)
	516	ND	...	...	...	...
	1819	SP	6090.0	-3.3 (0.8)	280 (40)	0.7 (0.2)
?	1881	SP	6088.8	-13 (4)	400 (50)	1.81 (0.3)
	2	SP (L)	6280.5	+0.3 (0.1)	280 (50)	1.4 (0.4)
	15	ND	...	...	...	...
	40	ND	...	...	...	...
	516	ND	...	...	...	...
	1819	ND	...	...	...	...
[O I] 6300	1881	ND	...	...	...	...
	2	ND	...	...	...	...
	15	SP	6302.1	-0.53 (0.02)	270 (20)	...
	40	SP	6303.8	-0.73 (0.05)	240 (60)	0.28 (0.05)
	516	SP	6302.5	-1.4 (0.1)	150 (20)	...
	1819	SP	6302.2	-7.6 (0.7)	260 (50)	0.84 (0.04)
[Fe X] 6374	1881	SP	6303.2	-9 (2)	400 (70)	1.0 (0.3)
	2	SP	6380.8	-0.7 (0.1)	700 (100)	1.2 (0.2)
	15	SP	6375.5	-0.8 (0.1)	230 (20)	...
	40	P-Cyg	6360.4	...	...	...
			6374.9	...	...	...
	516	ND	...	...	...	...
H $\alpha$	1819	SP	6377.3	-18 (3)	270 (20)	2.7 (0.1)
	1881	SP	6377.8	-26 (4)	430 (40)	3.9 (0.2)
	2	SP	6566.9	-105 (3)	556 (6)	660 (1)
	15	TP	6564.9, 6565.0, 6568.4	-302 (2)	69.6 (0.4), 494 (3), 1074 (4)	...
	40	SP (L)	6567.7	-680 (30)	490 (10)	292.7 (0.2)
	516	DP	6566.7	-260 (50)	140 (4), 280 (20)	...
He I 6678	585	SP	6566.7	-150 (30)	470 (20)	9.5(0.3)
	1819	SP	6566.4	-980 (50)	360 (8)	164.0 (0.3)
	1881	DP	0 6563.3, 0 6565.8	-1300 (100)	1700 (200), 570 (10)	256.7 (0.5)
	2	SP	6684.0	-3.98 (0.02)	580 (10)	17.4(0.9)
	15	SP	6681.2	-4.5 (0.2)	350 (10)	...
	40	SP	6682.7	-5.9 (0.2)	440 (20)	1.5 (0.2)
He I 7065	516	SP	6681.3	-7.2 (0.7)	210 (10)	...
	1819	SP	6681.0	-5 (1)	260 (50)	0.8 (0.1)
	1881	SP	6680.3	-1.2 (0.7)	440 (90)	1.3 (0.2)
	2	SP	7070.5	-5.44 (0.05)	500 (10)	30.4 (0.4)
	15	DP (LG)	7067.2,	-5.4 (0.8)	28 (3),	...

**Table 9**  
(Continued)

Line	Epoch	Profile	Line Center (Å)	EW (Å)	FWHM (km s <sup>-1</sup> )	Flux (10 <sup>-16</sup> erg s <sup>-1</sup> cm <sup>-2</sup> )
			7068.4		40 (20)	
	40	SP	7070.0	-9.8 (0.2)	428 (9)	4.46 (0.07)
	516	SP	7069.0	-6.6 (0.5)	200 (10)	...
	1819	SP	7067.7	-10 (1)	250 (20)	0.8 (0.1)
	1881	SP	7068.1	-18 (2)	280 (20)	1.5 (0.3)
He I 7281	2	ND	...	...	...	...
	15	SP	7284.8	-1.0 (0.1)	270 (40)	...
	40	ND	...	...	...	...
	516	ND	...	...	...	...
	1819	ND	...	...	...	...
	1881	ND	...	...	...	...
[CaII] 7291	2	ND	...	...	...	...
	15	SP	7293.7	>-0.2	≲16	...
	40	ND	...	...	...	...
	516	SP	7293.3	-0.5 (0.1)	≲40	...
	1819	ND	...	...	...	...
	1881	ND	...	...	...	...
[CaII] 7323	2	ND	...	...	...	...
	15	SP	7326.1	-0.123 (0.003)	≲20	...
	40	ND	...	...	...	...
	516	SP	7326.0	-0.9 (0.1)	≲60	...
	1819	ND	...	...	...	...
	1881	ND	...	...	...	...
O I 7774	2	P-Cyg	7765.4, 7783.1	...	...	...
	15	P-Cyg	7769.8, 7780.4	...	...	...
	40	P-Cyg	7771.7, 7781.0	...	...	...
	516	SP	7779.4	-1.9 (0.2)	90 (30)	...
	1819	ND	...	...	...	...
	1881	ND	...	...	...	...
[Fe XI] 7892	2	ND	...	...	...	...
	15	ND	...	...	...	...
	40	ND	...	...	...	...
	516	ND	...	...	...	...
	1819	SP	7895.3	-6 (2)	230 (40)	1.0 (0.2)
	1881	SP	7895.0	-17 (2)	400 (40)	1.9 (0.5)
O I 8446	2	SP (L)	8451.4	-5.6 (0.2)	160 (10)	6 (1)
	15	DP	8449.1, 8451.3	-10.3 (0.9)	35 (2), 470 (20)	...
	40	SP (L)	8452.1	-30 (10)	190 (20)	1.3 (0.2)
	516	SP	8450.0	-7 (2)	270 (20)	...
	585	SP	8450.3	-7 (1)	270 (60)	0.4 (0.1)
	1819	SP	8449.8	-6 (1)	270 (60)	6.3 (0.7)
	1881	SP	8449.4	-24 (4)	190 (10)	6.5 (0.8)
Ca II and Pa 8500	2	P-Cyg	8495.6, 8509.1	...	...	...
	15	DP	8500.6, 8503.7	-11.1 (0.6)	36 (1), 418 (7)	...
	40	SP	8503.6	-26 (1)	346 (9)	9.0 (0.1)
	516	DP	8504.3, 8500.8	-16 (1)	60 (5), 30 (2)	...
	585	SP	8502.3	-9 (3)	180 (30)	1.0 (0.2)
	1819	SP	8501.4	-7 (1)	620 (140)	1.4 (0.1)
	1881	SP	8498.5	-12 (2)	510 (60)	2.03 (0.09)
Ca II and Pa 8544	2	SP	8549.0	-2.30 (0.06)	420 (20)	19.0 (0.9)
	15	DP	8544.7, 8547.7	-9 (2)	39 (1), 520 (70)	...
	40	SP	8547.7	-20 (2)	320 (8)	7.56 (0.05)
	516	DP	8545.6, 8546.6	-16 (1)	43 (3), 270 (20)	...
	585	SP	8547.7	-7 (3)	240 (50)	0.7 (0.1)

**Table 9**  
(Continued)

Line	Epoch	Profile	Line Center (Å)	EW (Å)	FWHM (km s <sup>-1</sup> )	Flux (10 <sup>-16</sup> erg s <sup>-1</sup> cm <sup>-2</sup> )
Ca II and Pa 8664	1819	SP	8547.3	-7 (1)	190 (40)	0.6 (0.1)
	1881	SP	8545.7	-7 (1)	250 (40)	1.21 (0.06)
	2	SP – flat	8669.9	-4.8 (0.1)	540 (20)	22.6 (0.5)
	15	DP	8664.9, 8668.0	-9.3 (0.9)	37 (4), 460 (40)	...
	40	SP	8668.2	-16.7 (0.8)	301 (9)	5.30 (0.08)
	516	DP	8665.5, 8668.2	-18 (3)	58 (4), 340 (30)	...
	585	SP	8667.0	-8 (2)	190 (30)	<0.4
	1819	SP	8666.0	-7 (2)	130 (40)	1.3 (0.3)
	1881	DP	8666.7, 8672.2	-11 (1)	380 (70), 400 (300)	1.4 (0.1)
	Pa 8753	2	SP	8755.0	-1.88 (0.06)	410 (60)
15		DP	8753.1, 8755.9	-2.7 (0.2)	32 (5), 160 (10)	...
40		SP	8760.0	-13.1 (0.8)	550 (20)	3.6 (0.1)
516		SP	8754.8	-3.7 (0.3)	150 (20)	...
1819		ND	...	...	...	...
1881		SP	8755.0	-6 (2)	180 (60)	1.0 (0.2)
Pa 8865	2	P-Cyg	8851.9, 8871.7	...	...	...
	15	DP	8865.4, 8867.8	-4.7 (0.6)	100 (10), 400 (30)	...
	40	SP	8868.0	-6 (1)	360 (50)	3.0 (0.3)
	516	SP	8867.4	-4.0 (0.5)	100 (10)	...
	1819	SP	8866.5	-1.1 (0.9)	100 (50)	2.0 (0.3)
	1881	SP	8865.6	-7.2 (0.9)	600 (100)	1.3 (0.2)
Pa 9017	2	SP	9024.3	-6.5 (0.5)	840 (70)	26.7 (0.9)
	15	SP	9020.5	-4.9 (0.5)	410 (50)	...
	40	SP	9021.6	-11 (1)	430 (30)	2.3 (0.1)
	516	SP	9019.0	-4.6 (0.5)	180 (20)	...
	1819	ND	...	...	...	...
	1881	SP	9018.3	-7 (1)	290 (50)	0.3 (0.1)
Pa 9232	2	SP	9234.1	-2.4 (0.2)	820 (240)	8 (1)
	15	DP	9231.7, 9235.2	-2.4 (0.2)	49 (3), 110 (10)	...
	40	DP	9227.1, 9236.3	-7 (1)	180 (20), 630 (80)	2.8 (0.2)
	516	SP	9233.0	-5.0 (0.7)	150 (20)	...
	1819	SP	9233.0	-8 (1)	150 (30)	0.40 (0.09)
	1881	SP	9232.1	-5 (1)	290 (50)	0.7 (0.3)
Paε 9548	2	SP	9557.1	-0.77 (0.02)	180 (40)	50 (1)
	40	SP	9553.1	-19 (4)	500 (60)	5.4 (0.2)
Paδ 10052	1881	SP	9548.8	-9 (2)	280 (60)	4.2 (0.5)
	2	SP	10057.1	-7.2 (0.2)	152 (7)	32 (2)
	40	SP	10051.7	-0.7 (0.9)	430 (70)	3 (1)
	1881	SP	10052.6	-10 (3)	230 (40)	3.7 (0.5)

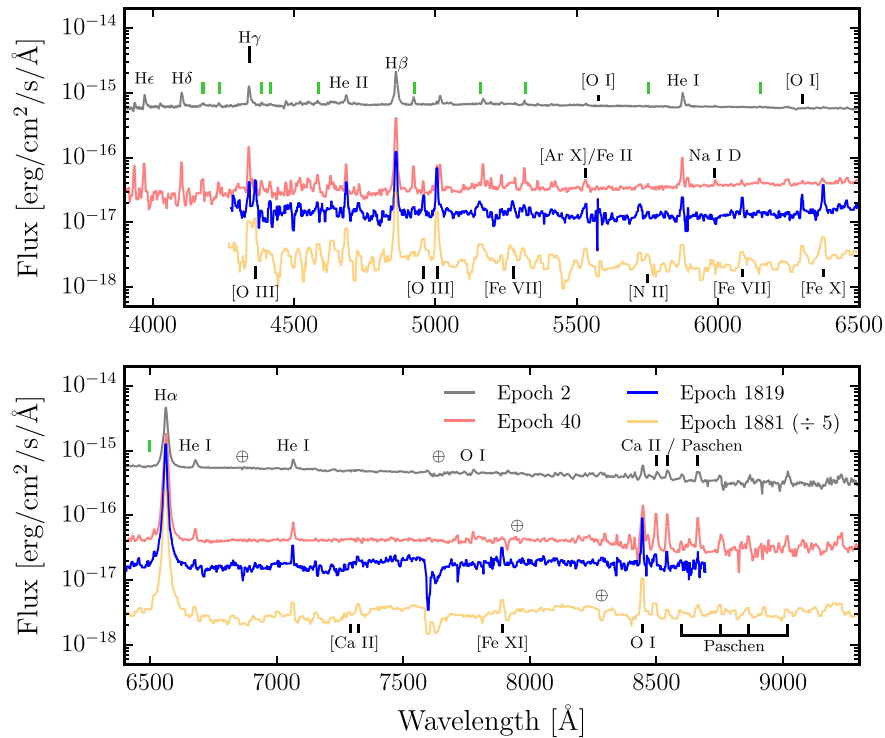
**Note.** Profiles are described as Gaussian (G), Lortenzian (L), Double Gaussian (DG), Gaussian and Lorentzian blend (LG), or P-Cygni profile (P-Cyg). Lines with missing epochs are not covered in the spectral range, while lines with ND are within the spectral range but not detected. Numbers in parentheses are  $1\sigma$  errors. The reported line center and FWHM errors do not account for uncertainties due to instrument resolution.

found in the eruptions of  $\eta$  Car ( $\approx 200$  days, from Binder et al. 2011). About 1450 days after the transient, we find an X-ray luminosity of  $\approx 2.4 \times 10^{36}$  erg s<sup>-1</sup>, which increases to  $\approx 5.9 \times 10^{37}$  erg s<sup>-1</sup> at about 1640 days. This increase in X-ray luminosity occurs at roughly the same time as the increase in optical/IR luminosity.

#### 4. DISCUSSION

Taken together, the X-ray, optical, and IR light curves and spectra consistently describe an HMXB undergoing an episode

of active accretion, which is fueled by persistent eruptions of the primary star, with SN 2010da representing the largest observed eruption in nearly a decade of observations. The optical and IR light curves are powered by the mass loss of the supergiant. This mass accretes onto the compact object, giving rise to X-ray emission. The X-rays in turn excite the high-ionization He II and coronal iron lines seen in the optical spectra. In this section, we summarize the properties of the progenitor and the surviving progeny of the 2010 eruption, and we discuss potential compact companions.



**Figure 8.** Low-medium-resolution optical spectra of SN 2010da. All spectra have been smoothed to a  $\approx 10$  Å resolution, and the 1881 day spectrum has been shifted downward by a factor of five for clarity. Unlabeled green lines refer to Fe II emission lines with a  $3\sigma$  detection in at least one epoch.

#### 4.1. The Progenitor of SN 2010da

Ignoring any contribution from a compact companion or accretion disk, our blackbody fit of the progenitor SED (with a temperature of 1500 K and a radius of 9.6 au) reveals a stellar bolometric luminosity of about  $2 \times 10^4 L_{\odot}$ . This luminosity is consistent with a  $15 M_{\odot}$  main-sequence star (Meynet & Maeder 2000) or a supergiant with a 10–12  $M_{\odot}$  ZAMS mass (Ekström et al. 2012). The low temperature suggests that this blackbody is not the photosphere of the progenitor. Instead, we interpret this SED as a dusty shell surrounding the star.

To further investigate progenitor candidates, we model the dusty environment of the progenitor and its SED using the radiative transfer code DUSTY (Ivezic et al. 1999). DUSTY is able to model the density profiles of spherically symmetric, radiatively driven winds, requiring as input the central source SED, the dust composition, the optical depth, and the inner dust temperature. Since we do not see silicate features around  $8 \mu\text{m}$  in our pre-eruption *Spitzer* observations, we choose a pure graphite environment (Drake & Ulrich 1980). The carbon-rich dust is consistent with the stability of the dust shell at a relatively high temperature ( $\approx 1500$  K), which has a higher sublimation temperature than silicate (Kobayashi et al. 2011). We assume that the shell has a thickness of  $R_{\text{out}}/R_{\text{in}} = 2$  and use a power-law density model that falls off as  $\rho \propto r^{-2}$ , typical of a wind. We additionally assume that the central source is a blackbody, and we leave its temperature as a free parameter. The final luminosity of the model is calculated using the normalized flux and radius computed by DUSTY. The UVOT observations during the 2010 outburst constrain the progenitor radius to be  $\approx 120 R_{\odot}$ . This limits the progenitor temperature to  $T \gtrsim 6200$  K. We are additionally unable to find satisfactory fits ( $\chi_r^2 < 2$ ) of the progenitor SED for temperatures above  $\approx 18,000$  K. The temperature of a  $15 M_{\odot}$  main-sequence star

is about 30,000 K, meaning that we can rule out such a progenitor. Due to the low luminosity, we can also rule out an LBV progenitor, which was previously suggested by others (Binder et al. 2016). The only remaining viable option at this luminosity is an evolved yellow or blue supergiant progenitor.

We can additionally use the DUSTY models to estimate the mass-loss rate of the progenitor. Following Kochanek et al. (2012), the mass loss is approximately equal to

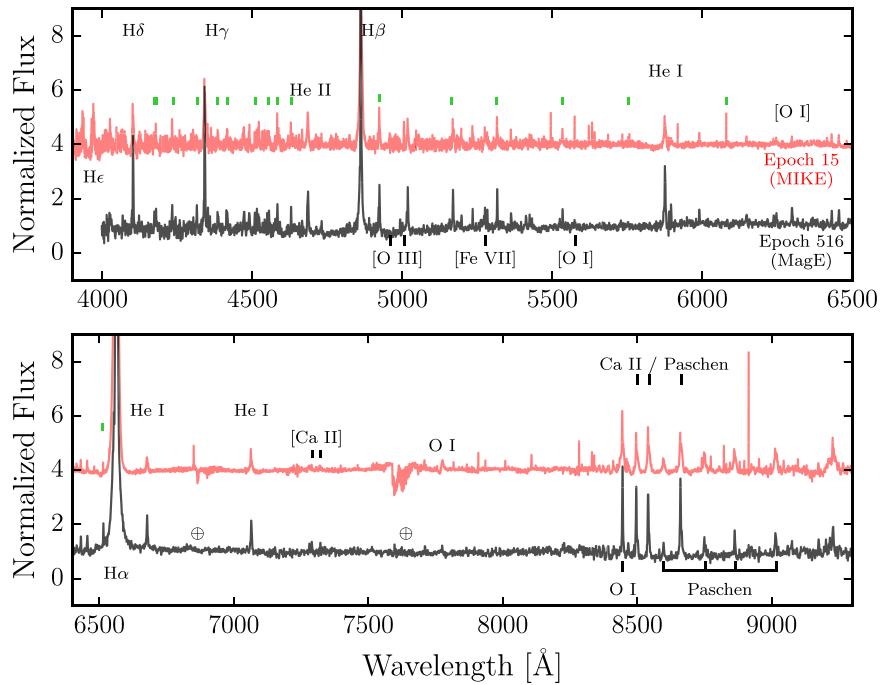
$$\dot{M} \approx \frac{\kappa_V}{8\pi v_w R_{\text{in}}}$$

where the opacity is  $\kappa_V \approx 120 \text{ cm}^2 \text{ g}^{-1}$ , we assume a wind velocity of  $v_w \approx 40 \text{ km s}^{-1}$  (the approximate line width of the narrow Balmer/He lines from the high-resolution MIKE spectrum), and  $R_{\text{in}}$  is the inner radius of the dusty shell as calculated by DUSTY. For the range of plausible models, the estimated mass-loss rates are  $(4\text{--}5) \times 10^{-7} M_{\odot} \text{ yr}^{-1}$ . This is in agreement with typical mass-loss rates of RSGs of this luminosity, significantly smaller than in super-AGB stars (Poelarends et al. 2008; Maun & Josselin 2011) and greater than in BSGs (Martins et al. 2015). However, asymmetry and inhomogeneity (e.g., clumpiness) in the CSM can greatly affect our estimated optical depth. A more extensive review of these effects can be found in Kochanek et al. (2012).

#### 4.2. The Progeny of SN 2010da and Its Environment

Our extensive photometric and spectroscopic data sets indicate that the source of SN 2010da is still active and underwent a dramatic transition to a bluer and hotter SED with a smaller radius of  $\approx 6$  au after the 2010 eruption. Additionally, the progeny is significantly more luminous than the progenitor by a factor of approximately two to five. Although it is possible





**Figure 9.** High-resolution optical spectra of SN 2010da from MIKE and MagE. Both spectra have been smoothed to  $\approx 4$  Å resolution for display purposes. Unlabeled green lines refer to Fe II emission lines with a  $3\sigma$  detection in at least one epoch.

that the bolometric luminosity of the progenitor was larger than we predict with a significant fraction of light contributed at longer wavelengths from cool dust, which was heated during the transient, it is most likely that the ongoing mass ejections and their interaction with a compact companion/CSM are injecting additional energy into the system.

In addition to being brighter, the source is also undergoing significant variability in the optical of  $\approx 1$ – $2$  mag within a few hundred days. The variability and bolometric magnitude of the progeny ( $M_{\text{bol}} \approx -7$ ) are reminiscent of supergiant long-period variables (Wood et al. 1983), though these do not typically show B[e]-like emission lines in their spectra nor are they often surrounded by a thick CSM.

To constrain the progeny properties, we use DUSTY to model the SEDs around 560 and 1880 days. Again using the constraint from the UVOT light curve, we find that the progeny is hotter than  $\approx 8900$  K. Additionally, at temperatures higher than  $\approx 25000$  K, the estimated radius becomes atypically small for a supergiant (i.e.,  $\lesssim 15 R_{\odot}$ ), though we can find acceptable fits beyond this temperature. To reiterate, we have previously ruled out a main-sequence star as the progenitor of SN 2010da, meaning that the progeny must also be an evolved supergiant. These temperature and luminosity constraints are shown in an HR diagram in Figure 19. We can again calculate the mass-loss rates at these different epochs, this time assuming that a new wind of  $v_w \approx 200$  km s $^{-1}$  has formed. We find a mass-loss rate of  $\dot{M} \approx 3 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$  at 560 days and a slightly larger rate of  $\dot{M} \approx 6 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$  at 1880 days. These numbers are consistent with the mass-loss rate before the outburst.

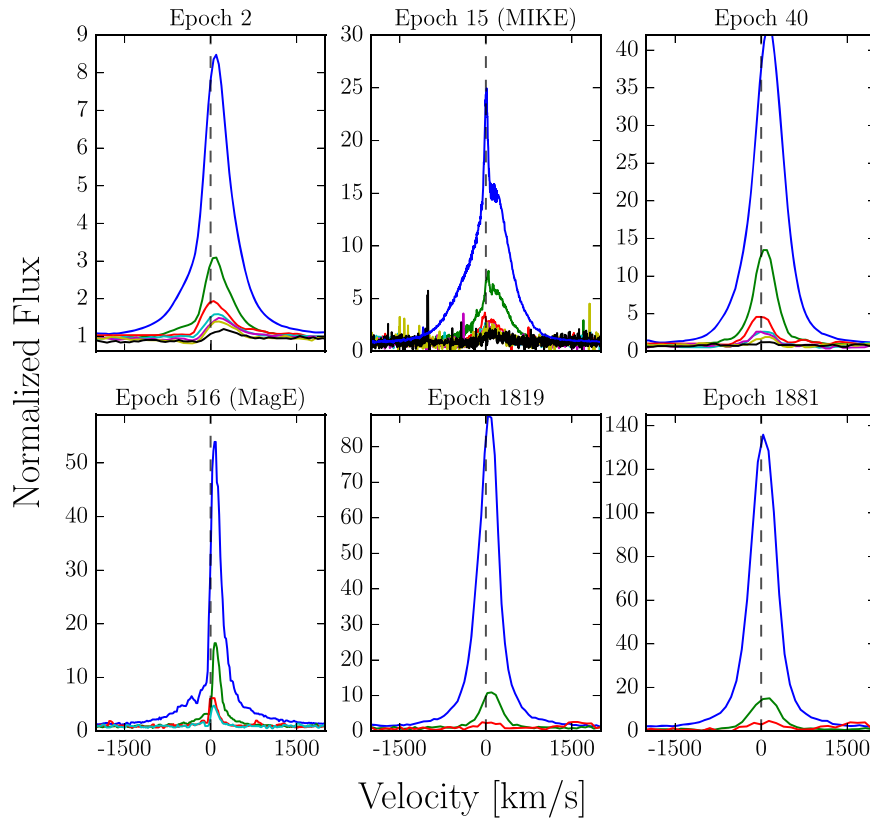
The low luminosity, B[e] features and bluer SED are all consistent with an RSG transitioning into a blue-loop phase of its evolution (Langer & Heger 1997). Additionally, the widths and shapes of the multi-component emission lines are consistent with a newly formed wind interacting with existing mass loss seen in the early stages of a blue-loop phase of an

RSG (Chita et al. 2008). The blue loop occurs when RSGs evolve off the Hayashi line toward the BSG regime as their envelope structure shifts from convective to radiative. During this transition, the envelope of the RSG will spin up and the radius will drastically decrease (Heger & Langer 1998). During this evolution, the star can reach its critical rotation rate and develop a slow equatorial outflow, leading to B[e]-like emission lines.

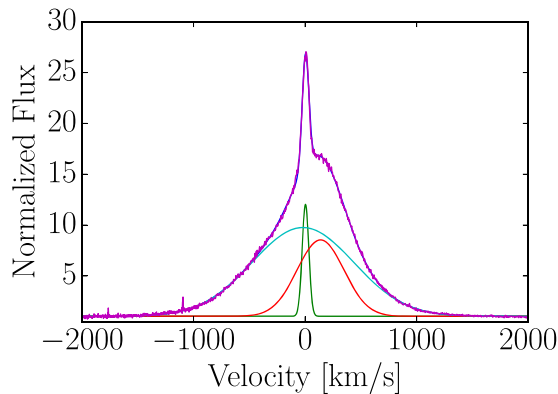
The environment surrounding the progenitor/progeny is extremely complex, as indicated by the varying estimated electron densities from the Balmer decrement and the existence of forbidden calcium and iron lines. Like many supergiants, SN 2010da might be surrounded by a clumpy wind, which can explain the low- and high-density regions necessary to excite the various emission lines detected in our spectra. The progenitor’s dust shell at  $\approx 10$  au seems to have been at least partially destroyed by the initial transient based on the strong initial UV and X-ray detections. However, the continued infrared excess and SED shape suggest that either some of this dust survived or new dust has since formed at  $\approx 6$  au. The surrounding CSM is carbon rich and irradiated by X-ray/UV emission from the compact binary companion, meaning some dust must be continuously destroyed and formed. During periods of eruptions and enhanced accretion, the UV emission excites coronal iron lines in the CSM, which we observe in the most recent optical spectra taken at 1819 and 1881 days.

#### 4.3. SN 2010da as a HMXB

Based on the strong X-ray luminosity ( $\sim 10^{37}$  erg s $^{-1}$ ) detected well before and after the optical transient, the strong He II 4686 Å emission, the coronal iron lines, and the hard X-ray spectrum, we conclude that SN 2010da is in a supergiant X-ray binary system exhibiting B[e] phenomena. A similar conclusion was reached by Binder et al. (2011) and Lau et al. (2016). However, it is difficult to make a definitive statement



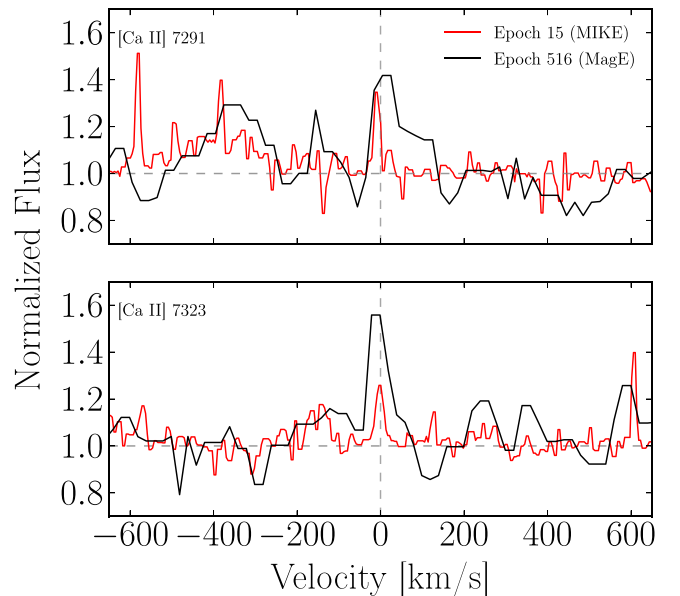
**Figure 10.** Evolution of Balmer lines in the optical spectra. Each epoch is listed above its corresponding subplot.  $H\alpha$ ,  $H\beta$ ,  $H\gamma$ ,  $H\delta$ ,  $H\epsilon$ ,  $H\zeta$ , and  $H\eta$  are shown in blue, green, red, teal, pink, yellow, and black, respectively. Note that the y-axes are all on independent scales for clarity. The local continuum has been normalized to one by fitting a first-order polynomial.



**Figure 11.**  $H\alpha$  emission line 14 days after the outburst. The line is well described by three Gaussian components (shown in cyan, red, and green). The local continuum has been normalized to one by fitting a first-order polynomial.

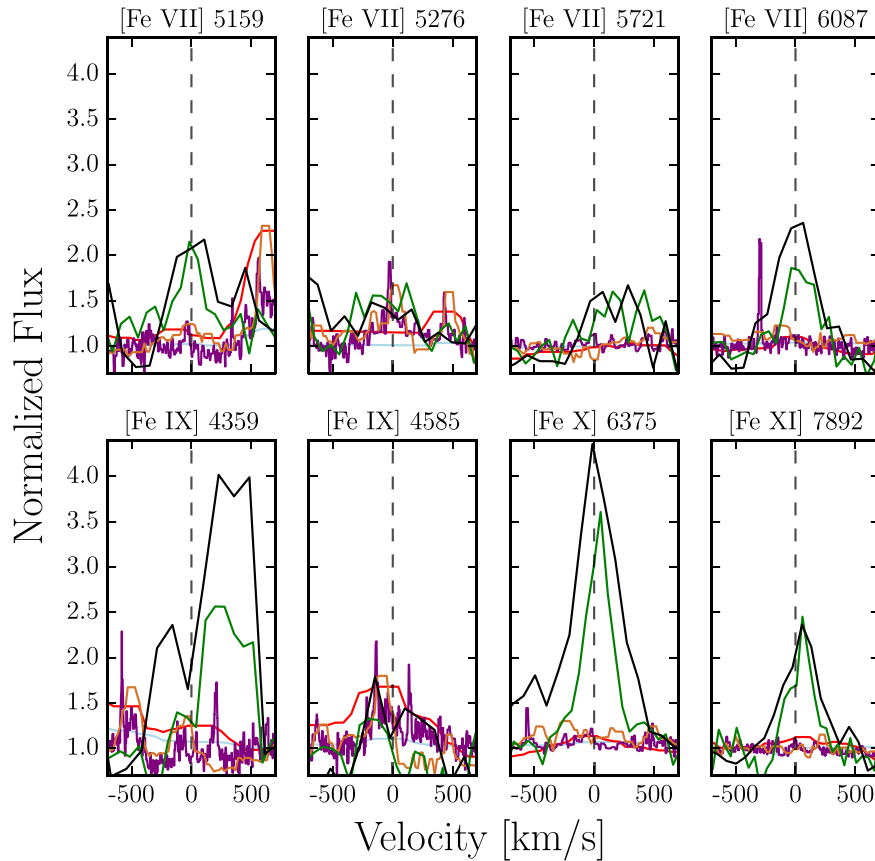
about the nature of the compact object itself. The ultraluminous X-ray transient is far above the Eddington limit of a  $1.4 M_{\odot}$  neutron star, but the hard spectrum ( $\Gamma \approx 0$ ) and the high X-ray luminosity are consistent with other SGXBs with neutron star companions, such as Vela X-1 (Lewin et al. 1997; Binder et al. 2011; Wang 2014). It is possible to explain the super-Eddington luminosity of the initial outburst by invoking beaming along the line of sight or large magnetic fields (Mushtukov et al. 2015). In fact, a ULX powered by a neutron star was recently discovered with an X-ray luminosity greater than the peak luminosity of SN 2010da (Bachetti et al. 2014).

SN 2010da also exhibits B[e] phenomena, consistent with a B[e] X-ray binary. Such binaries typically undergo two types of

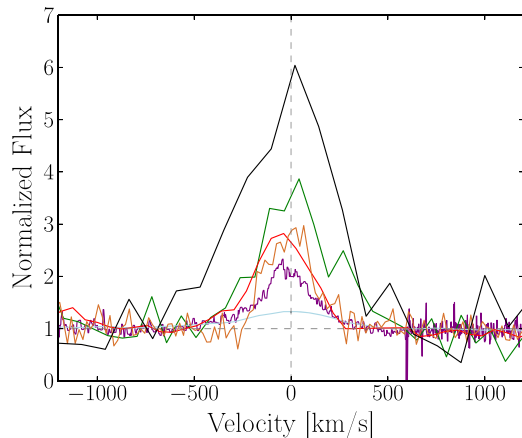


**Figure 12.** Narrow, [Ca II] lines detected in our high-resolution spectra. The local continuum has been normalized to one by fitting a first-order polynomial.

transients: dimmer ( $L_X \sim 10^{36-37}$  erg  $s^{-1}$ ), shorter ( $\tau \sim$  days) Type I outbursts, which are associated with the orbital period of the binary, and brighter ( $L_X \gtrsim 10^{37}$  erg  $s^{-1}$ ), longer ( $\tau \gtrsim$  weeks) Type II outbursts, which are possibly associated with the disruption of the B[e] disk (Reig 2011). The disk-disruption theory has undergone recent criticism following the discovery of several disks that have remained intact after a Type II

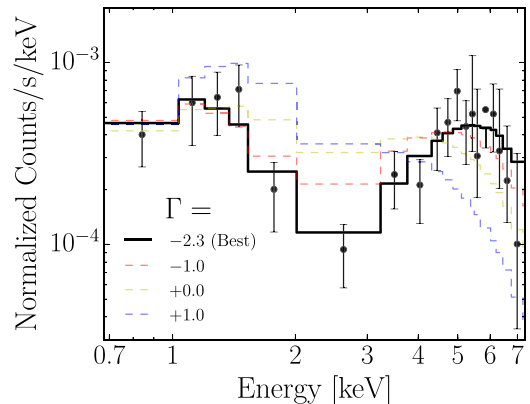


**Figure 13.** Evolution of the coronal iron lines. Spectra taken on days 2, 15, 40, 516, 1819, and 1881 are shown in light blue, purple, red, orange, green, and black, respectively. The local continuum has been normalized to one by fitting a first-order polynomial.



**Figure 14.** Evolution of the He II 4686 Å emission. Spectra taken on days 2, 15, 40, 516, 1819, and 1881 are shown in light blue, purple, red, orange, green, and black, respectively. The local continuum has been normalized to one by fitting a first-order polynomial.

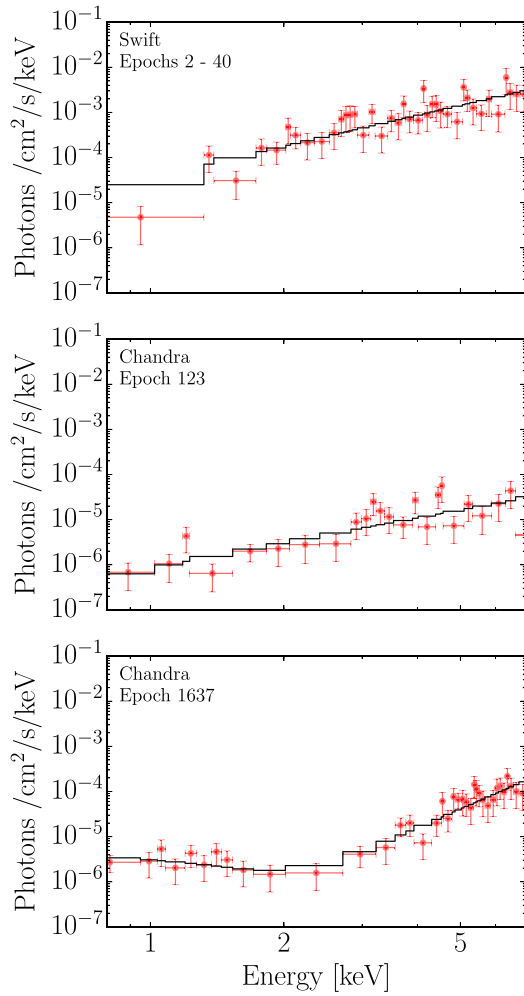
outburst (Reig & Fabregat 2015). The duration ( $\sim 50$  days) and hard spectral index ( $\Gamma \sim 0$ ) of the progeny of SN 2010da are consistent with a Type II outburst (Reig & Nespoli 2013). However, the X-ray luminosity during the transient ( $L_X \approx 6 \times 10^{39} \text{ erg s}^{-1}$ ) is much more luminous than typical Type II outbursts ( $L_X \approx 10^{37} - 10^{38} \text{ erg s}^{-1}$ ). Because little is known about the physical origin of Type II outbursts, we cannot definitely say if SN 2010da is an unusual Type II outburst or a new type of X-ray transient associated with eruptive stellar mass loss.



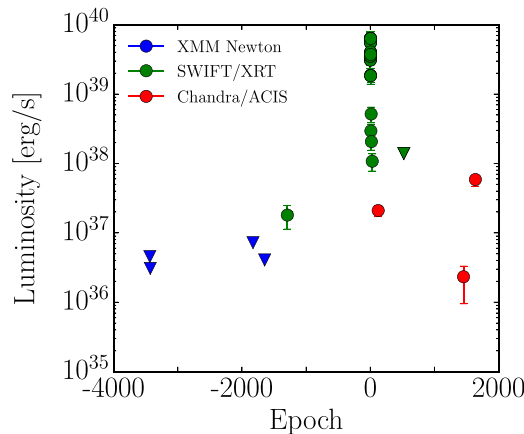
**Figure 15.** Latest X-ray spectrum of SN 2010da from *Chandra* Observation 16029, normalized by the detector’s effective area. Shown are four models with a Bremsstrahlung and power-law component. We fix the photon index,  $\Gamma$ , to  $-1.0$  (red),  $0.0$  (yellow), and  $+1.0$  (blue) and compare to the best-fit model with  $\Gamma = -2.3$ . We are unable to recover the bimodal structure of the spectrum with softer power laws.

#### 4.4. Comparison to Other Dusty ILOTs and Impostors

Although the canonical model of dusty ILOTs are massive LBVs ejecting dense shells of mass, it has become clear in recent years that these events arise from a variety of progenitors (Berger et al. 2009; Smith et al. 2011; Kochanek et al. 2012). Most of the well-studied ILOTs and their progenitors lie in one of two observational classes. The first class is made up of objects with blue and luminous progenitors, such as LBVs or yellow hypergiants (e.g., SN 2009ip, SN 1954J). ILOTs in the

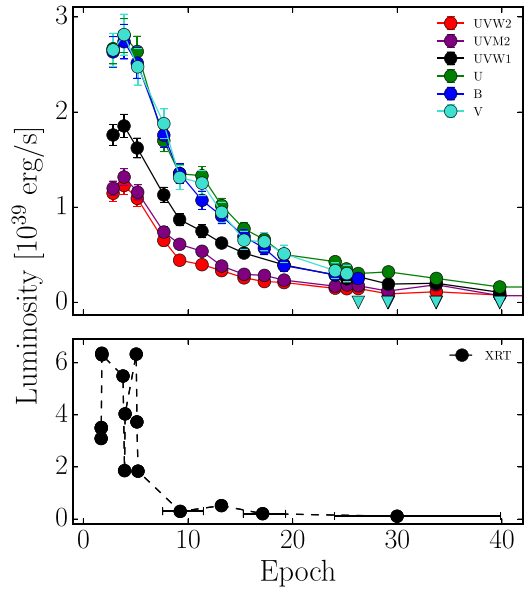


**Figure 16.** Top: the *Swift* X-ray spectrum created using data taken within 40 days of the transient. The best-fit power-law model is overlaid. Middle/bottom: the *Chandra* X-ray spectra (Obs. ID 12238 and 16029) with best-fit power-law and disk models (see the text for details).

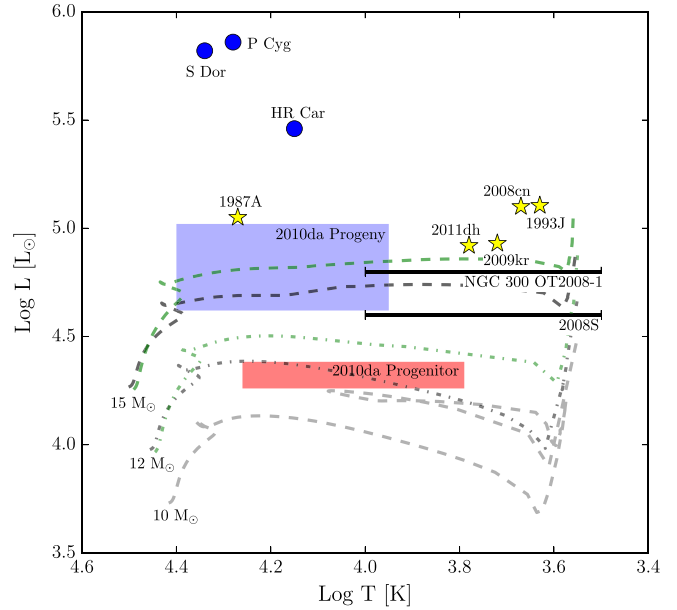


**Figure 17.** X-ray light curve of SN 2010da. Triangles indicate  $3\sigma$  upper limits. Downward facing triangles are  $3\sigma$  upper limits.

“blue” class survive their transients and can undergo multiple eruptions. Objects in this class include rare  $\eta$  Carinae analogs such as the recent UGC 2773-OT (Smith et al. 2016) and several ILOTs associated with yellow hypergiants undergoing LBV-like outbursts, like SN Hunt 248 (Mauerhan et al. 2015)



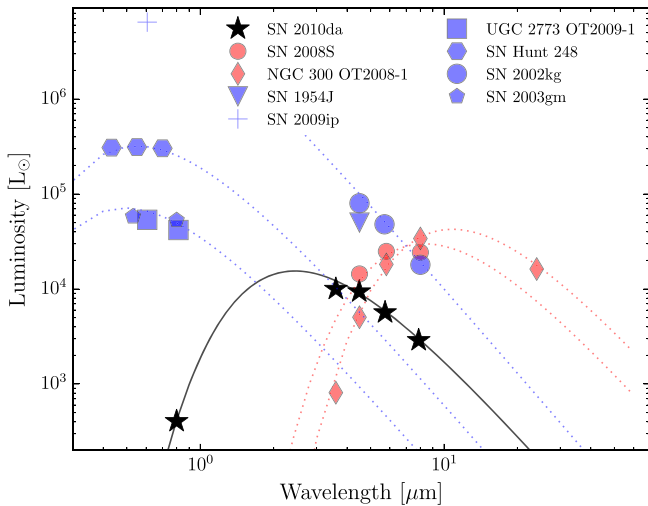
**Figure 18.** XRT (top) and UVOT light curves near the time of discovery. Downward facing triangles are  $3\sigma$  upper limits.



**Figure 19.** Stellar evolutionary tracks of  $10 M_{\odot}$ ,  $12 M_{\odot}$ , and  $15 M_{\odot}$  RSG models with (green) and without (black) rotation from Ekström et al. (2012) compared to the estimated temperatures and luminosities of progenitor (red) and progeny (blue) of SN 2010da. For comparison, we also show the progenitors of NGC 300 OT2008-1 (Prieto 2008), SN 2008S (Prieto et al. 2008), several supernovae and three well-studied LBVs: S Doradus, P Cygni, and HR Carinae (Humphreys et al. 2011).

and PSN J09132750+7627410 (Tartaglia et al. 2016). The second class of ILOTs is made up of objects with red and extremely cool ( $T \sim 100$  s K) progenitor SEDs. These ILOTs appear to be terminal explosions, which are potentially electron-capture SNe from massive AGB stars (e.g., SN 2008S, NGC 300 OT2008-1), though other theories exist to explain these events (Smith et al. 2011; Kochanek et al. 2012; Adams et al. 2016).

Does the system hosting SN 2010da fit into one of these two classes? We directly compare the progenitor, transient, and progeny associated with SN 2010da to two red dusty ILOTs



**Figure 20.** Progenitor SEDs of LBV-like (blue) and SN 2008S-like (red) ILOTs compared to the SED of the progenitor of SN 2010da (black). SN 2010da sits between these two classes. The dashed lines are blackbody fits for select objects to guide the eye.

(SN 2008S and NGC 300 OT2008-1) and two blue ILOTs thought to be LBVs (SN 1954J, or Variable 12 in NGC 2403, and SN 2009ip).

SN 2008S and NGC 300 OT2008-1 had peak absolute magnitudes of  $M_V \approx -14$  and  $M_V \approx -12$ , respectively. These two objects exhibited similar properties and have since faded beyond their initial progenitor luminosities in the IR (Adams et al. 2016). SN 2008S and NGC 300 OT2008-1 had progenitors whose SEDs were consistent with cool circumstellar dust ( $T \approx 300\text{--}500$  K) and large radii ( $R \approx 150\text{--}350$  au; Prieto 2008; Khan et al. 2010). These temperatures are about four times cooler than the progenitor of SN 2010da ( $\approx 1500$  K), and their estimated radii are about 10 times larger. The luminosities of these progenitors were approximately two to three times higher than the progenitor of SN 2010da. On the opposite end of the ILOT spectrum lie the blue ILOTs: SN 1954J, a massive star in NGC 2403, which underwent an LBV-like eruption and remains active today, and SN 2009ip, an LBV in NGC 7259, which likely exploded in 2012 (Smith et al. 2010; Margutti et al. 2013; Mauerhan et al. 2013). Prior to 1949, the progenitor of SN 1954J had a blue magnitude of  $M_b \approx -6.6$  (assuming a distance modulus of 27.6; Smith et al. 2001). Similarly, the progenitor of SN 2009ip was extremely bright ( $M_{\text{Bol}} \approx -10$ ) and variable by as much as one magnitude before its 2009 outburst. Both progenitors of these blue ILOTs are notably brighter and bluer than the progenitor of SN 2010da. The progenitor SED of SN 2010da sits between these two classes, as shown in Figure 20.

These objects show similar diversity during their transient light curves. Within the first month of discovery, the red ILOTs (SN 2008S and NGC 300 OT2008-1) experienced a similar decay rate of  $\approx 0.03$  mag  $\text{d}^{-1}$  (Berger et al. 2009)—much more slowly than SN 2010da, which decayed at  $\approx 0.1$  mag  $\text{d}^{-1}$ . Although NGC 300 OT2008-1’s light curve steepens at later times (to  $\approx 0.06$  mag  $\text{d}^{-1}$ ), it does not exceed the decline rate of SN 2010da. In contrast, the decline rate of SN 2009ip’s 2009 outburst within the first month ( $\approx 0.2$  mag  $\text{d}^{-1}$ ) is faster than that of SN 2010da (Smith et al. 2010). In the case of SN 2009ip, such a fast decline rate was attributed by Smith et al.

(2010) to the ejection of an optically thick shell, which is not ruled out as a possibility for SN 2010da.

Spectroscopically, SN 2010da shares features with both the red and blue ILOT classes. For example, the red ILOTs and SN 2010da share similar narrow Balmer and forbidden calcium lines, with  $H\alpha$  reaching a maximum width of  $\approx 1200$  km  $\text{s}^{-1}$ . Like NGC 300 OT2008-1, we detect He I emission in SN 2010da, but we additionally detect He II due to the X-ray/UV-enriched environment from the compact companion. Most notably, unlike NGC 300 OT2008-1, our high-resolution spectrum reveals Balmer lines that are weakly asymmetric and lacking any absorption; high-resolution spectra of NGC 300 OT2008-1 reveal  $H\alpha$  emission with clear absorption slightly blueward of rest wavelength (Berger et al. 2009; Bond et al. 2009). Similarly, the blue ILOTs are also dominated by hydrogen Balmer and Fe II emission (typical of hot LBVs) with  $\text{FWHM} \approx 550$  km  $\text{s}^{-1}$  (Smith et al. 2010; Margutti et al. 2013). Unlike SN 2010da, there was no [Ca II] emission detected in SN 2009ip, though [Ca II] emission has been detected in eruptions of cool LBVs such as UGC 2773-OT (Smith et al. 2010). Late-time spectra of SN 1954J reveal broad  $H\alpha$  emission with  $\approx 700$  km  $\text{s}^{-1}$ , broader than what is observed in the progeny of SN 2010da.

One of the most notable differences between SN 2010da/the red ILOTs and the blue ILOTs is the fate of their progeny. The blue ILOTs underwent clearly non-terminal eruptions (excluding the 2012 explosion of SN 2009ip; Margutti et al. 2013; Mauerhan et al. 2013). Specifically, recent photometry shows that the progeny of SN 1954J has since faded by  $\approx 2$  mag in the optical and is now consistent with a blackbody with a temperature of  $\approx 6500$  K. This has been interpreted as an  $\eta$  Car analog which is now shrouded in a dusty nebula similar to  $\eta$  Car’s Homunculus (Smith et al. 2001; Van Dyk et al. 2005). The most recent SED of SN 1954J is much bluer than that of SN 2010da and suggests a notably higher bolometric luminosity ( $\approx 10^5 L_\odot$ ). Based on luminosity and the SED, SN 2010da is unlikely to be an LBV outburst. In contrast, the progenies of SN 2008S and NGC 300 OT2008-1 have faded past their progenitors in the IR, leading some authors to argue that they were electron-capture supernovae from super AGB stars (Botticella et al. 2009; Thompson et al. 2009; Adams et al. 2016). The clear re-brightening of the progeny of SN 2010da several hundred days after the 2010 eruption illustrates that it is not a member of this red class of transients, but its similarities might point to a related progenitor that is entering the last phase of its life.

Thus, SN 2010da is unlike many of the previously studied ILOTs. First, the transient is not energetic enough to be a true LBV outburst. We can roughly estimate the energy radiatively emitted from SN 2010da as  $\approx L_{\text{peak}} t_{1.5}$ , where  $L_{\text{peak}}$  is the peak luminosity and  $t_{1.5}$  is the time it takes the transient to dim by 1.5 magnitudes (see Smith et al. 2011). We estimate  $t_{1.5} \lesssim 30$  days based on the upper limit reported by Monard (2010), and we estimate the peak luminosity to be  $L_{\text{peak}} = 4.5 \times 10^{39}$  erg  $\text{s}^{-1}$ . The total radiative energy is thus  $\lesssim 10^{46}$  erg. This is less energetic than the typical LBV outburst ( $\approx 10^{47}$  erg; Smith et al. 2011). SN 2010da is also less energetic than the red SN 2008S-like ILOTs, which radiate about  $L_{\text{peak}} t_{1.5} \approx 5 \times 10^{47}$  erg. Additionally, SN 2008S-like events are either terminal or produce progeny that are notably dimmer than their progenitors (Adams et al. 2016); the progeny is currently more luminous than its progenitor by a factor of  $\approx 5$ .

## 5. SUMMARY AND CONCLUSIONS

We presented comprehensive, multiwavelength observations (X-ray, UV, optical, and IR) of the dusty ILOT SN 2010da, extending thousands of days before and after the outburst. These observations allowed us to study the progenitor, outburst, and progeny in great detail. Due to its low luminosity and red SED, SN 2010da seems inconsistent with an LBV outburst as interpreted by Binder et al. (2016). From our data set, we conclude that SN 2010da was the eruption of a massive star ( $\sim 10\text{--}12 M_{\odot}$ ) exhibiting B[e] phenomena. The high-resolution spectra exhibit double-peaked Balmer and Paschen emission lines with narrow components consistent with a pre-existing RSG wind and a newly formed supergiant wind. This suggests that the star responsible for SN 2010da may be a YSG transitioning onto a blue loop. The optical variability and iron/calcium emission indicate a complicated CSM, which is repeatedly disturbed by mass loss of the primary star.

The supergiant responsible for SN 2010da is likely the primary star of an HMXB. The system shows consistently high X-ray luminosity ( $L_X \approx 10^{37} \text{ erg s}^{-1}$ ), and during the 2010 event, the system underwent an ultraluminous X-ray outburst ( $L_X \approx 6 \times 10^{39} \text{ erg s}^{-1}$ ). Late-time emission of coronal iron lines are fueled by a hot, X-ray- and UV-rich region near this binary. While we cannot make a definitive statement about the nature of the compact object, dedicated and deep X-ray observations may shed light on its nature.

SN 2010da is unique in the heterogeneous class of ILOTs. Its progenitor was dimmer and bluer than the AGB-like progenitors of dusty ILOTs NGC 300 OT2008-1 and SN 2008S; however, it is notably dimmer and redder than LBVs and yellow hypergiants experiencing similar outbursts. Also unlike other dusty ILOTs and supernova impostors, the progeny of SN 2010da is more luminous than its progenitor in both the IR and optical. The progeny is still undergoing significant outbursts, and continued follow-up is crucial to understanding the elusive nature of this object.

Like many ILOTs, SN 2010da marks an important point in stellar evolution of increased activity and mass loss. SN 2010da highlights the diversity of dusty ILOTs and the need for multiwavelength photometric and high-resolution spectral follow-up to understand these objects. There is no doubt that future facilities such as LSST will populate the intermediate luminosity gap that currently exists. Extensive spectroscopic follow-up of current events will allow us to identify archetypes, like SN 2010da, of classes that will arise from these surveys.

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