

A large population of strongly lensed faint submillimetre galaxies in future dark energy surveys inferred from *JWST* imaging

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ABSTRACT

Bright galaxies at sub-millimetre wavelengths from *Herschel* are now well known to be predominantly strongly gravitationally lensed. The same models that successfully predicted this strongly lensed population also predict about one percent of faint 450 μm -selected galaxies from deep JCMT surveys will also be strongly lensed. Follow-up ALMA campaigns have so far found one potential lens candidate, but without clear compelling evidence e.g. from lensing arcs. Here we report the discovery of a compelling gravitational lens system confirming the lensing population predictions, with a $z_s = 3.4 \pm 0.4$ submm source lensed by a $z_{\text{spec}} = 0.360$ foreground galaxy within the COSMOS field, identified through public *JWST* imaging of a 450 μm source in the SCUBA-2 Ultra Deep Imaging EAO Survey (STUDIES) catalogue. These systems will typically be well within the detectable range of future wide-field surveys such as *Euclid* and *Roman*, and since sub-millimetre galaxies are predominantly very red at optical/near-infrared wavelengths, they will tend to appear in near-infrared channels only. Extrapolating to the *Euclid*-Wide survey, we predict tens of thousands of strongly lensed near-infrared galaxies. This will be transformative for the study of dusty star-forming galaxies at cosmic noon, but will be a contaminant population in searches for strongly lensed ultra-high-redshift galaxies in *Euclid* and *Roman*.

Key words: gravitational lensing: strong – galaxies: evolution – infrared: galaxies – submillimetre: galaxies

1 INTRODUCTION

Strong galaxy-galaxy gravitational lensing is the phenomenon in which light from a background galaxy (source) is bent around a foreground galaxy (lens) due to gravity, resulting in a distorted view of the source reaching an observer, typically in the form of extended arcs, rings, and multiple images of the source. Observations of these objects can provide a wealth of information, from testing general relativity and dark matter models by constraining the mass profile of the foreground galaxy dark matter halo (Miralda-Escude & Babul

1995; Sonnenfeld et al. 2015; Collett et al. 2018) and detecting substructures within the halo and along the line of sight (Mahdavi et al. 2007; Vegetti et al. 2014; Hezaveh et al. 2016; Amorisco et al. 2022; He et al. 2022; Ostdiek et al. 2022; Wagner-Carena et al. 2023), to helping understand galaxy formation and evolution by studying the inner mass profile slope of the lens (Cui et al. 2017; Holanda et al. 2017; Sonnenfeld & Cautun 2021) and the properties of high-redshift background sources (Swinbank et al. 2010; Danielson et al. 2011; Dye et al. 2015; Geach et al. 2018; Salmon et al. 2018; Collett & Smith 2020; Khullar et al. 2021). Specific types of strong lenses can also aid in constraining cosmological parameters, with the matter density Ω_m and dark energy equation of state parameters w_0 and w_a

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constrained by double source plane lenses (Collett et al. 2012; Smith & Collett 2021; Sharma & Linder 2022), while time delays between different images of the same source can put constraints on the Hubble constant H_0 from lensed quasars (Treu 2010; Denzel et al. 2021; Li et al. 2021) and supernovae (Collett et al. 2019; Bag et al. 2021; Bayer et al. 2021).

These applications require large statistical samples of lenses. Following earlier work on the fraction of lensed submillimetre galaxies (SMGs) in blank field SMG surveys (Chapman et al. 2002), an early breakthrough of the *Herschel* space observatory was the observation that nearly all the bright SMGs (500 μm flux densities above ~ 100 mJy) are strongly gravitationally lensed, once obvious contaminant populations of blazars and very nearby galaxies are removed (Negrello et al. 2010). This enabled a strong lens selection method with an extraordinary $\sim 100\%$ efficiency. Further campaigns at longer wavelengths with data from the South Pole Telescope found a similarly efficient strong lens selection (Reuter et al. 2020). Since SMGs are typically very red at optical and near-infrared wavelengths (e.g. Chen et al. 2016; Michałowski et al. 2017; Lim et al. 2020; Shim et al. 2022), follow-up redshift campaigns have mostly relied on CO and [CI] transitions, yielding several hundred secure redshifts to date (e.g. Neri et al. 2020; Reuter et al. 2020; Urquhart et al. 2022), which in turn have enabled a very wide range of galaxy evolution studies exploiting the strong gravitational magnification (e.g. Zavala et al. 2018; Bendo et al. 2023a; Hagimoto et al. 2023; Reuter et al. 2023).

The search for strong gravitational lenses is on the cusp of another transformative change. Upcoming surveys such as *Euclid* (Laureijs et al. 2011) and the Vera Rubin Legacy Survey of Space and Time (LSST; Ivezić et al. 2019) are set to discover around 10^5 of them amongst a larger population of $\sim 10^9$ galaxies in the optical and near-infrared, with near-infrared imaging of approximately 17000 strong lenses (Weiner et al. 2020) also expected by the future *Nancy Grace Roman Space Telescope* (Green et al. 2012; Spergel et al. 2015; Akeson et al. 2019).

Most of the effort to date in making lensing predictions for these future dark energy surveys, and in designing machine learning algorithms to identify the lensed systems, has focused on the higher angular resolution visible wavelength imaging (e.g. Laureijs et al. 2011; Collett 2015; Metcalf et al. 2019). In this work, we show that the relatively neglected near-infrared channels will also be a very rich source of strong gravitational lenses. We do this by using new public *JWST* imaging of the 450 μm SCUBA-2 SMGs in the COSMOS field to search for strong gravitational lensing features, demonstrating that the faint SMG population has a very high surface density of strong lensing systems visible in the near-infrared.

This paper is organised as follows. The catalogue of SCUBA-2 SMGs in the COSMOS field and the corresponding *JWST* data are discussed in Section 2. The results of the investigation are presented in Section 3, including photometry and modelling of the identified gravitational lens and its impact on the inferred number density of strongly-lensed SMGs. This is followed by a discussion of our findings in Section 4 and final conclusions in Section 5.

2 DATA

2.1 Submm Observations

The 450 μm imaging used in this work was taken by the Submillimetre Common-User Bolometer Array 2 (SCUBA-2; Chapin et al. 2013; Dempsey et al. 2013; Holland et al. 2013) on the James Clerk Maxwell Telescope (JCMT) as part of the SCUBA-2 Ultra Deep

Imaging EAO Survey (STUDIES; Wang et al. 2017) within the 2 square degree Cosmological Evolution Survey (COSMOS; Scoville et al. 2007) field. Centred on (ra, dec) coordinates (10:00:28.0, +02:24:00) (Casey et al. 2013; Lim et al. 2020), STUDIES observations covered 450 square arcminutes down to a depth of 0.59 mJy per beam at 450 μm with the resulting data set containing 479 identified SMGs (Gao et al., in prep.).

2.2 Infrared Observations

The *JWST* data consisted of publicly available Near Infrared Camera (NIRCam) and Mid-Infrared Instrument (MIRI) observations of the COSMOS field as part of the COSMOS-Web (Casey et al. 2022) and Public Release IMAGING for Extragalactic Research (PRIMER; Dunlop et al. 2021) surveys. Within the COSMOS field, PRIMER in part aims to cover 144 square arcmin in 8 NIRCam filters and 112 square arcmin in two MIRI filters. Likewise, COSMOS-Web aims to cover 0.54 square degrees in four NIRCam filters down to 5σ depths of ~ 28 magnitudes, and 0.19 square degrees in one MIRI filter down to 5σ depths of ~ 26 magnitudes.

Data was acquired through the Space Telescope Science Institute (STScI) Mikulski Archive for Space Telescopes (MAST) Portal, covering the STUDIES area using a search radius of 15 arcmin centred on the STUDIES central coordinates following Lim et al. (2020). As the COSMOS-Web and PRIMER surveys were not yet completed, only 59 per cent of the search area had been surveyed at the time of this work. All available COSMOS-Web and PRIMER calibration level 3 imaging data within the search area was acquired: COSMOS-Web data was partially available for all filters (NIRCam F115W, F150W, F277W, and F444W along with MIRI F770W); PRIMER data was partially available for eight out of ten filters (NIRCam F115W, F200W, F277W, F356W, F410M, and F444W along with MIRI F1800W and F770W; missing NIRCam F090W and F150W).

3 RESULTS

3.1 Photometry

At present, out of the 479 SMGs identified at 450 μm by STUDIES, 283 were found to have coordinates within regions imaged by *JWST* in at least one filter: 63 SMGs were imaged in COSMOS-Web and 263 in PRIMER. From visually inspecting the corresponding *JWST* NIRCam and MIRI imaging, a single strong lens candidate was discovered, whose measured properties from STUDIES are given in Table 1. The foreground lens of this lensing system corresponded to the known galaxy WISEA J100023.99+021749.6, at (ra, dec) coordinates (10:00:24.025, +02:17:49.54) with a spectroscopic redshift of 0.35978 ± 0.00018 (Damjanov et al. 2018), which has not as yet been classed as a gravitational lens in previous studies. The SMG itself has submm colours clearly strongly inconsistent with the foreground lens redshift, as the 450 μm : 850 μm (deboosted) flux density ratio of ~ 1.9 (Table 1) requires a redshift of $z > 3$ (e.g. Lim et al. 2020), indicative of a background source. Imaging at the SMG's position had only been taken in *JWST* MIRI filters F770W and F1800W, with the central foreground lens and lensed background source seen in both, as shown in Figure 1. The arc of the lensed background source is clearly visible in the higher resolution F770W filter, and two separate objects are still resolved in the F1800W filter despite its lower resolution.

For cross-matching, the background source in PRIMER imaging is located at 10:00:23.98 +02:17:50.26, while the STUDIES 450 μm

Table 1. STUDIES SCUBA-2 data for the 450 μm SMG. Columns are as follows: object name; right ascension (J2000) (deg); declination (J2000) (deg); signal-to-noise ratio; raw flux density (mJy); instrumental noise at the 450 μm SMG position (mJy); deboosted flux density (mJy); total noise (consisting of instrumental noise, flux-deboosting uncertainty, and confusion noise) at 450 μm (mJy); deboosted flux density of the SMG at 850 μm (mJy); and the total noise at 850 μm (mJy).

ID	RA (deg)	Dec (deg)	SNR	$S_{450,\text{raw}}$	$\delta S_{450,\text{raw}}$	$S_{450,\text{total}}$	$\delta S_{450,\text{total}}$	$S_{850,\text{total}}$	$\delta S_{850,\text{total}}$
STUDIES-COSMOS-450-019 /	150.100117	2.297500	15.79	19.2	1.2	18.5	2.3	9.76	0.65
STUDIES-COSMOS-850-006				mJy	mJy	mJy	mJy	mJy	mJy

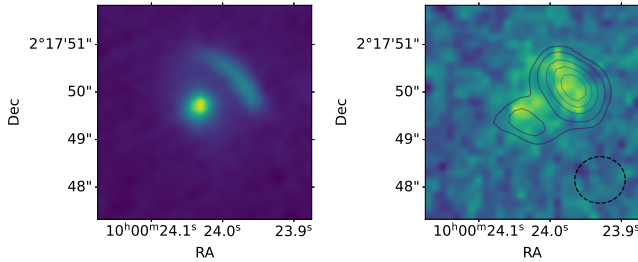


Figure 1. *JWST* MIRI F770W (left) and F1800W (right) imaging of the strong gravitational lens from PRIMER as 4.5×4.5 arcsec cutouts centred on the visible foreground lens. The lensed background source lies just to its upper right, clearly visible as an arc in the left image. ALMA imaging is presented as overlaid contours on the right image, depicting the lensed background source as a bright image and fainter counterimage, with no foreground lens visible. The ALMA beam FWHM is shown as the dashed black line in the right image. Cutouts have been reprojected to align their sky orientation with figure axes, inline with ALMA observations.

source is located 1.054 arcsec away at 10:00:24.03 +02:17:51.00. The point-spread function (PSF) of the SCUBA-2 beam has a full width at half maximum (FWHM) of 7.9 arcsec for 450 μm (Dempsey et al. 2013; Lim et al. 2020), hence neighbouring galaxies near the lensing system could not be ruled out as candidates for the SMG from this data alone. However, publicly available high-resolution Atacama Large Millimeter/submillimeter Array (ALMA) submm observations from the ALMA public archive confirmed the lensed background source to be the SMG. This source was found to be the bright source AS2COS0005.1 observed in the 870 μm AS2COSMOS catalogue of the brightest SCUBA-2 SMGs in COSMOS (Simpson et al. 2019, 2020), with a likely counter image detected close to the foreground lens given by the nearby source AS2COS0005.2, which is 4.0 ± 0.7 times fainter than the main extended image. These 870 μm ALMA observations of AS2COS0005.1/5.2 are shown here as contours overlaying the PRIMER *JWST* F1800W image in Figure 1.

PRIMER *JWST* photometry is given in Table 2 for the foreground lens and background source separately. As there existed no *JWST* F1800W photometry for the lensing system, the F1800W flux densities were calculated based on the same lens and source segmentation applied to the F770W photometry, with uncertainties based on the 1σ background. Photometry at other wavelengths exists for only a single object: that of the lens, the source, or the combined lensing system, depending on the wavelength. Table 3 contains the publicly available flux densities of the lensing system in other wavelengths or passbands, with those shorter than 5 μm and those longer than 100 μm corresponding to the lens and the source, respectively.

To obtain a photometric redshift of the lensed source, the 450 μm and 850 μm flux densities in Table 1 along with the *Herschel* SPIRE and MAMBO photometry in Table 3 were used to fit a spectral energy distribution (SED) using the ‘‘Eyelash’’ template, which was used to

fit the $z > 2$ SMG of the Cosmic Eyelash gravitational lens (see e.g., Swinbank et al. 2010; Iglesias-Groth et al. 2017; Duivenvoorden et al. 2018), resulting in a redshift of $z_{\text{source}} = 3.41 \pm 0.14$. However, SMGs exhibit a range of intrinsic colour temperatures (e.g. Blain et al. 2003), resulting in larger photometric redshift uncertainties (e.g. Casey 2020). Recently, Bendo et al. (2023b) compared spectroscopic redshifts in bright lensed *Herschel* galaxies with photo- z from the Eyelash template, using submm and mm-wave data from ALMA and SPIRE; the diversity of SEDs in that sample implies an irreducible uncertainty in our photo- z of ± 0.4 , so we adopt $z_{\text{source}} = 3.4 \pm 0.4$.

To verify that the photometry indicated a combination of lens and source, the flux densities in Table 3 up to wavelengths of 4.6 μm were then used to fit an SED for the lens at its much lower redshift, taken to be a 5 Gyr-old dustless early-type galaxy (initial mass function from Chabrier 2003, metallicity $\log_{10}(Z/Z_{\odot}) = 0.0$), using the Flexible Stellar Population Synthesis for Python package PYTHONFSPS (Conroy et al. 2009; Conroy & Gunn 2010)¹. The resulting SEDs and photometry are presented in Figure 2, which provides a good fit to the overlaid *JWST* PRIMER photometry. To explore the robustness of the lens SED fitting, a more flexible SED modelling was also performed, fitting a single stellar population with free star formation history, dust extinction, and metallicity. This provided an almost identical fit, and resulted in the following parameters for the foreground lens: $\log_{10}(\text{age/yr}) = 9.8$, extinction $A_V = 0$, star formation rate $\log_{10}(\text{SFR}) = -1.43$, and stellar mass $\log_{10}(M_*/M_{\odot}) = 10.69$, which are consistent with an early-type galaxy. The stellar mass also agrees with the value obtained for the foreground lens from the hCOSMOS redshift survey (Damjanov et al. 2018), which used the MMT Hectospec multifibre spectrograph to obtain a velocity dispersion $\sigma = 153 \pm 25 \text{ km s}^{-1}$ and stellar mass $\log_{10}(M_*/M_{\odot}) = 10.55^{+0.14}_{-0.11}$.

While a limit could not be placed on the stellar mass of the background source due to the effects of dust reddening, the star formation rate of the background source could be estimated, for which we used the latest *Planck* cosmology (Aghanim et al. 2020). Integrating the fitted Eyelash SED presented in Figure 2 over the mid and far infrared (8-1000 μm) produced a luminosity of $L_{\text{FIR}} = 6.8 \times 10^{12} \mu^{-1} L_{\odot}$, where μ is the magnification factor. Using the conversion $\text{SFR}(M_{\odot} \text{ yr}^{-1}) = 4.5 \times 10^{-44} L_{\text{FIR}}(\text{ergs s}^{-1})$ from Kennicutt (1998), results in a star formation rate of $\sim 1200 \mu^{-1} M_{\odot} \text{ yr}^{-1}$ for a Salpeter initial mass function (IMF). The inferred SFR would be a factor $\times 1.5$ lower for a Kroupa IMF (see Hayward et al. 2014, also for other caveats on this calculation).

Public *Hubble Space Telescope* (*HST*) WFC3 imaging of the lens was also available from the Cosmic Assembly Near-IR Deep Extragalactic Legacy Survey (CANDELS). However, while the foreground lens was clearly visible across all available bands, there was no obvious background source at these wavelengths. For example, modelling and subtraction of the foreground lens light profile for the near-infrared F160W image of the system was performed using

¹ <https://doi.org/10.5281/zenodo.591505>

Table 2. PRIMER *JWST* MIRI F770W and F1800W data for the foreground lens and background source. Columns are as follows: object; right ascension (J2000) (deg); declination (J2000) (deg); F770W and F1800W flux densities (mJy); F770W and F1800W AB magnitudes; and F770W ellipticity and position angle counter-clockwise from north (deg). F1800W flux densities were not available from *JWST* photometry, so were calculated based on the same lens and source segmentation used for F770W photometry, with uncertainties based on the $1-\sigma$ background.

Object	RA (deg)	Dec (deg)	$F_{\nu, F770W}$ (mJy)	$F_{\nu, F1800W}$ (mJy)	$m_{AB, F770W}$	$m_{AB, F1800W}$	e	ϕ (deg)
lens	150.100134	2.297140	$(17.59 \pm 0.07) \times 10^{-3}$	$(6.06 \pm 2.52) \times 10^{-3}$	20.787 ± 0.004	$21.95^{+0.58}_{-0.38}$	0.082	-17.06
source	150.099914	2.297295	$(9.81 \pm 0.06) \times 10^{-3}$	$(7.87 \pm 2.77) \times 10^{-3}$	21.421 ± 0.007	$21.66^{+0.47}_{-0.33}$	0.675	42.98

Table 3. Combined photometry for the lensing system, excluding *JWST* and SCUBA-2 data, presented in Figure 2. Source references: (1) Adelman-McCarthy et al. (2008) (2) Darvish et al. (2015) (3) Cutri et al. (2013) (4) Ashby et al. (2015) (5) Spitzer Science Center (SSC) & Infrared Science Archive (IRSA) (2021) (6) Casey et al. (2012) (7) Simpson et al. (2020) (8) Bertoldi et al. (2007). SDSS data are obtained via best-fitting de Vaucouleurs and exponential light models. VISTA K_s uncertainty is based on the 5σ depth of the UltraVISTA near-infrared imaging survey (McCracken et al. 2012). *Spitzer* IRAC data are corrected to total flux from a 2.4" aperture.

Wavelength / Observed Passband	F_{ν} (mJy)	Ref.
4858 Å (SDSS g)	0.0039 ± 0.0005	(1)
6290 Å (SDSS r)	0.0191 ± 0.0009	(1)
7706 Å (SDSS i)	0.0309 ± 0.0014	(1)
9222 Å (SDSS z)	0.0446 ± 0.0061	(1)
2.2 μm (VISTA K_s)	0.121 ± 0.001	(2)
3.4 μm (WISE W1)	0.068 ± 0.006	(3)
3.6 μm (<i>Spitzer</i> IRAC)	0.062 ± 0.004	(4)
4.5 μm (<i>Spitzer</i> IRAC)	0.056 ± 0.004	(4)
4.6 μm (WISE W2)	0.048 ± 0.011	(3)
12.0 μm (WISE W3)	< 0.302	(3)
22.0 μm (WISE W4)	< 2.98	(3)
24.0 μm (<i>Spitzer</i> MIPS)	< 0.0328	(5)
250 μm (<i>Herschel</i> SPIRE)	16.5 ± 2.6	(6)
350 μm (<i>Herschel</i> SPIRE)	25.5 ± 4.5	(6)
500 μm (<i>Herschel</i> SPIRE)	22.6 ± 6.4	(6)
870 μm (ALMA)	$8.4^{+0.4}_{-0.3}$	(7)
1.20 mm (MAMBO)	5.2 ± 0.9	(8)

GALFIT, fitting the lens light with a Sérsic component and an exponential disk component convolved with the *HST* F160W model PSF, with no visible background source detected down to a 2σ AB magnitude depth of 25.5, as shown in Figure 3. For comparison, *Euclid* has an average 5σ limiting AB magnitude of 24.5 in the near-infrared (*Euclid* Collaboration et al. 2022), so for this lens it would also be unable to detect the background source at this wavelength, but the SMG is both unusually faint in the optical and at unusually high redshift for this flux density (see Section 4).

3.2 Lens Modelling

Lens modelling was performed on the F770W image with LENSTRONOMY² (Birrer & Amara 2018; Birrer et al. 2021) using a singular isothermal ellipsoid (SIE) lens mass model, along with elliptical Sérsic profiles for the light models of the lens and source convolved with the *JWST* model PSF for MIRI F770W. Lens modelling was also performed for an elliptical power law (EPL) mass profile as well as modelling with and without external shear, but these offered no better fit than an SIE alone. Due to limited resolution, the mass

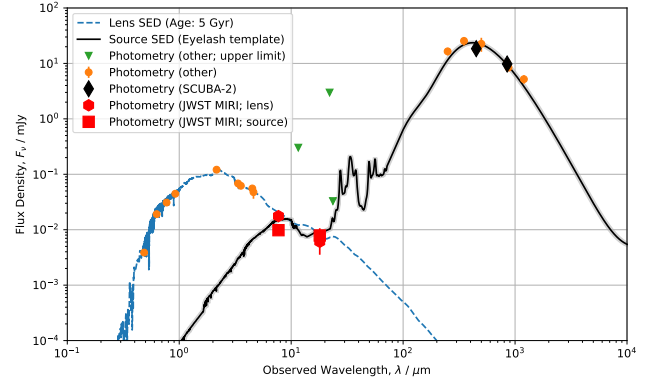


Figure 2. Spectral energy distributions (SEDs) of the foreground lens (blue dashed line) and background source (black solid line), with the shaded region around the latter giving the 1σ errors. The lens SED is of a 5 Gyr-old dustless early-type galaxy described in the text at the observed spectroscopic redshift of $z_{\text{lens}} = 0.35978$, fitted to photometry from other surveys (orange circles; see Table 3) primarily associated with the lens (those with wavelengths shorter than $\sim 5 \mu\text{m}$). The source SED is the “Eyelash” template used to fit the Cosmic Eyelash gravitational lens, which gave a photometric redshift of $z_{\text{source}} = 3.4 \pm 0.4$ when fitted to the STUDIES SCUBA-2 photometry (black diamonds; see Table 1) as well as the *Herschel* SPIRE and MAMBO photometry. Also displayed are upper limits of photometry from other surveys (green downward triangles; see Table 3) as well as the *JWST* MIRI F770W and F1800W photometry for the lens (red hexagon) and source (red square) presented in Table 2.

models produced very similar images despite converging on different parameters, indicating a degeneracy in the modelling. However, the use of a power law or external shear resulted in higher chi-squared errors and excessively elliptical lens mass profiles compared to the lower-ellipticity lens light, indicating that the SIE alone provided the best fit.

The resulting SIE lens mass model parameters are given in Table 4, with Figure 4 presenting the reconstructed lens and source, the residuals between the observed and modelled lens, and the convergence and magnification maps from the modelling. We follow the definition of ellipticity widely used in lens modelling: for axis ratio $q = b/a$, where a and b are the major and minor axes respectively, and position angle ϕ of the major axis, complex ellipticity components are given by

$$e_1 = \frac{1-q}{1+q} \cos 2\phi \quad e_2 = \frac{1-q}{1+q} \sin 2\phi \quad (1)$$

such that the parameter space (e_1, e_2) can be used for more effective fitting.

The modelled lens light has an effective radius $R_{e, \text{lens}} = 0.61 \pm 0.03$ arcsec, a Sérsic index $n_{\text{lens}} = 5.0^{+0.1}_{-0.2}$ consistent with an early-type galaxy, and an axis ratio $q_{\text{light}} = 0.73 \pm 0.02$ closely matching its modelled mass profile. The resulting fit reconstructs the source

² <https://github.com/lenstronomy/lenstronomy>

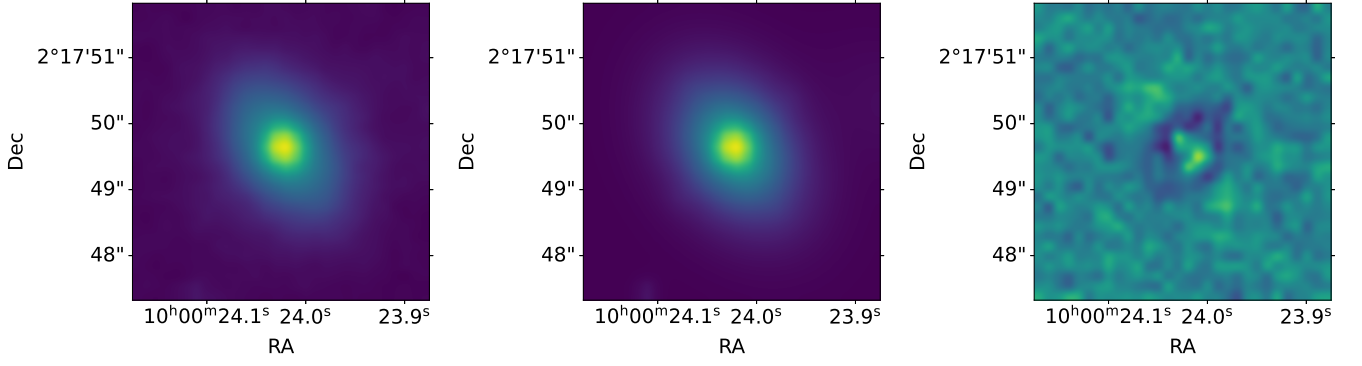


Figure 3. GALFIT lens light modelling and subtraction for publicly-available *Hubble Space Telescope* (*HST*) F160W near-infrared imaging. The lens light was fitted with a Sérsic component and an exponential disk component convolved with a F160W model PSF. From left to right: the observed cutout image; the GALFIT model for the lens light; and the residual image. All cutouts share the same angular scale and sky orientation matching those of Figure 1, with the first and second likewise presented with logarithmic scaling.

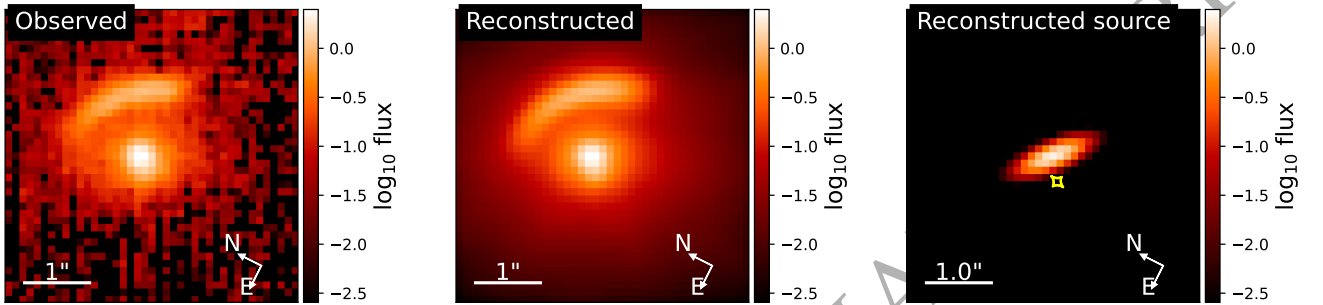


Figure 4. Results from lens modelling. From left to right: the observed cutout image, with original *JWST* sky orientation used for the modelling; the reconstructed image from modelling; and the reconstructed source image centred on the source with the lens model caustic overlaid in yellow. All cutouts share the same angular scale and sky orientation, and are presented with logarithmic scaling.

Table 4. Modelling parameters of the lens mass profile fitted to the *JWST* F770W image, which is centred on the (ra, dec) coordinates (10:00:24.025, +02:17:49.54) of WISEA J100023.99+021749.6. The lens mass was fitted to a singular isothermal ellipsoid (SIE) model with the values given below for each parameter: Einstein radius (arcsec); the two components of complex ellipticity e_1 and e_2 ; the corresponding axis ratio q and position angle ϕ (deg) clockwise from east; the central RA & Dec offset of the mass profile relative to the cutout centre (arcsec); and the total magnification of the lensed source μ . The light profiles of both the lens and source were modelled with elliptical Sérsic profiles.

θ_E (arcsec)	e_1	e_2	q	ϕ (deg)	RA offset (arcsec)	Dec offset (arcsec)	μ
$0.63^{+0.01}_{-0.01}$	$-0.11^{+0.04}_{-0.04}$	$-0.07^{+0.03}_{-0.03}$	$0.77^{+0.08}_{-0.07}$	106^{+11}_{-9}	$0.10^{+0.01}_{-0.01}$	$0.15^{+0.01}_{-0.01}$	$3.0^{+0.3}_{-0.3}$

image with an effective radius $R_{e,\text{source}} = 0.17 \pm 0.01$ arcsec, and a low Sérsic index $n_{\text{source}} = 0.50^{+0.02}_{-0.01}$ consistent with a late-type galaxy. The unlensed source is also highly elliptical, with an axis ratio $q_{\text{source}} = 0.34 \pm 0.02$, which matches the general results of previous samples of SMGs (e.g. [Chen et al. 2022b](#); [Cheng et al. 2022](#)). The source’s major axis aligns with the direction of elongation of the lensed arc, resulting in a lower magnification than expected: from the ratio of the lensed and unlensed model source flux above the 1σ background, a magnification of $\mu = 3.0 \pm 0.3$ is required to convert the highly elliptical source into the observed extended arc. The modelling also uncovered a faint point-source counterimage within a region of the lens light similar to the counterimage observed with ALMA.

Lens modelling was also performed on the ALMA data in order to constrain the SMG source properties at longer wavelengths. For this, the lens model was constrained to the parameters in Table 4 of

the *JWST* F770W lens model, while the source parameters were left unconstrained and initialised at those predicted by the F770W source model. A magnification of $\mu = 3.0^{+1.5}_{-1.2}$ was required to produce the ALMA source’s lensed image and counterimage, in agreement with the magnification obtained from modelling the F770W data, with the modelling limited by comparatively low resolution and a large restoring beam FWHM. However, the modelled (ra, dec) position of the unlensed source was found to be offset by $(0.03 \pm 0.03, -0.11 \pm 0.02)$ arcsec relative to that of the source at F770W. This suggests that two different regions of the SMG are responsible for the emission seen in *JWST* F770W and ALMA. Such a \sim kpc-scale offset has precedents in other strongly gravitationally lensed SMGs (e.g. [Dye et al. 2015](#)).

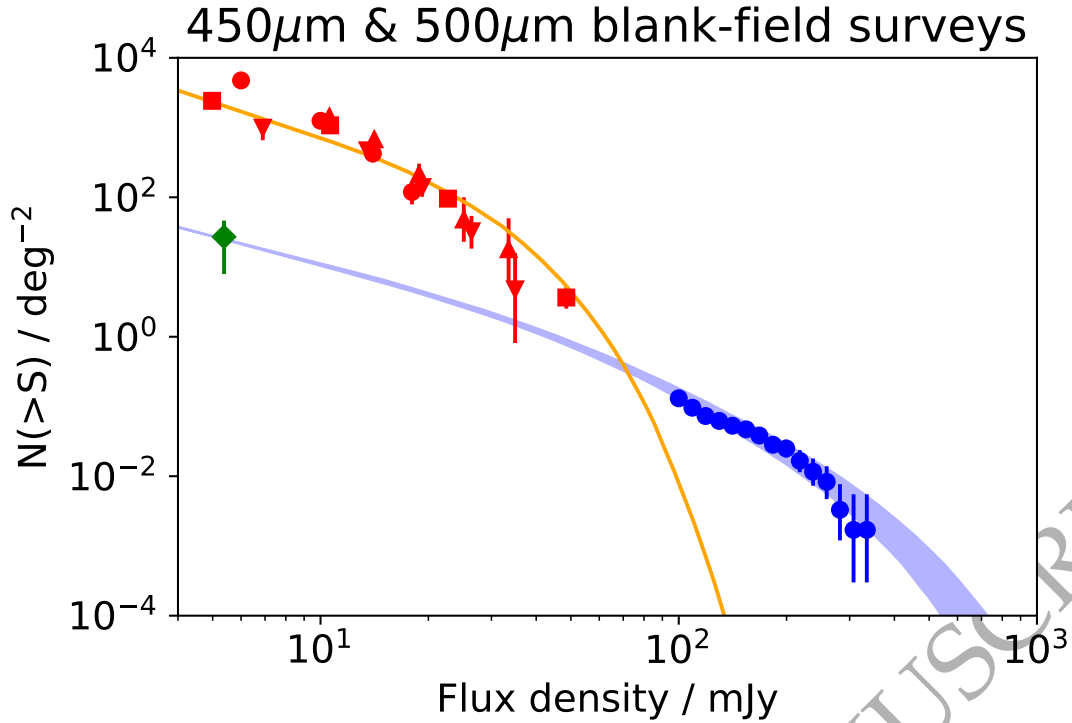


Figure 5. Integral blank-field galaxy number counts at 450-500 μm . The submm-bright strongly gravitationally lensed population of Negrello et al. (2017, 500 μm) is shown as blue filled circles. Our inferred 450 μm surface density of submm-faint strongly lensed galaxies is shown as a green diamond, calculated based on the $2 \pm \sqrt{2}$ lenses found within the 59% of the 450 square arcminute effective area of STUDIES that has been mapped by *JWST* at the time of this work. For this, we use the median 450 μm flux density of the STUDIES catalogue (Gao et al., in prep.) that was searched to obtain the lens candidates. The unlensed 450 μm -selected populations are shown in red, as follows: Zavala et al. (2017), red circles; Chen et al. (2013), downward-pointing red triangles; Casey et al. (2013), upward-pointing red triangles. The unlensed 500 μm -selected population of Glenn et al. (2010) is also shown as red squares. The lensed and unlensed population models are shown as blue and orange curves, respectively (Negrello et al. 2017; Trombetti et al. 2021, Negrello, priv. comm.), with the blue shaded area showing the lensed models when considering maximum magnifications in the range 10-15 experienced by the background source population as a whole. The models are calculated for the *Herschel* SPIRE 500 μm filter, and are identical for the SCUBA-2 450 μm filter (see e.g., Geach et al. 2013) as the former’s bandwidth encompasses that of the latter. Despite the extrapolation of over an order of magnitude fainter than the *Herschel* lensed galaxy number counts, our inferred source density agrees extremely well with the model predictions.

3.3 Inferred Number Density of Strongly-Lensed SMGs

The strongly-lensed SMGs found within the area of the COSMOS field can be used to better predict the number density of these objects across the sky (e.g. Chapman et al. 2002), and hence how large a population may be visible in the near-infrared by telescopes like *Euclid* and *Roman*. In addition to this lens, there is another SMG strong lens candidate known to be present in the survey area, albeit without a distinctive arc morphology: ALMA follow-ups of SMGs in the COSMOS field (Chen et al. 2022a) identify the galaxy AS2COSMOS0002.1 as being strongly lensed. They found the SMG to have a spectroscopic redshift of $z = 4.5956 \pm 0.0006$ and a 450 μm flux density of 22 ± 2 mJy. They also estimated the magnification to be $\mu = 3.0^{+1.4}_{-0.7}$, derived from the presence of several $z \sim 1$ foreground galaxies, suggestive of a galaxy group lens.

In Figure 5, we compare our inferred surface density of strongly gravitationally lensed galaxies with the model predictions presented in Trombetti et al. (2021). The lensed models (blue) fit the bright lensed *Herschel* number counts from Negrello et al. (2017, blue circles) very well; our inferred lensing number count of 27 ± 19 per square degree (green diamond) based on the two lenses identified within the *JWST*-mapped STUDIES area agrees strikingly well with the model prediction, despite the extrapolation to fainter flux den-

ties of a factor of $\sim \times 10$. For comparison we also show the model unlensed number counts (orange line), strongly constrained by the fit to the observed unlensed number counts (red data points). For the Casey et al. (2013, upward triangles) and Zavala et al. (2017, red circles) surveys, we report the integral counts in those studies, while for the Chen et al. (2013, downward triangles) survey we have recalculated the integral counts using the reported differential number counts, differing slightly from the methodology adopted in Negrello et al. (2017) and Trombetti et al. (2021). We incorporate the asymmetry of the reported uncertainties using the procedure³ in Appendix B of Laursen et al. (2019). We also show the integral counts implicit from the $P(D)$ analysis of Glenn et al. (2010, red squares), using the same methodology to incorporate asymmetric uncertainties; we base our estimates on their differential number count constraints using a cosmic far-infrared background prior from the Far Infrared Absolute Spectrophotometer (FIRAS). The unlensed model agrees strikingly well with the FIRAS-constrained data and most of the other blank-field survey data, with discrepancies attributable to large-scale structure.

³ Implementation in Python at https://github.com/anisotropela/add_asym

The flux densities for integral number counts are taken to be the faintest end of the faintest bin of the corresponding differential number counts used for that measurement. For our data point we use the median 450 μm flux density of the STUDIES catalogue (Gao et al., in prep.), rather than the minimum SMG flux density or noise limit, as the nature of the SCUBA-2 observations results in the noise limit systematically increasing from the beam centre to the outskirts, resulting in some SMGs with flux densities that reach below the noise levels experienced by many others. Combined with our limited sample size of lenses, we hence use the median flux density, which is also almost identical to the median 3.5σ background used as the threshold for SMG identification.

4 DISCUSSION

Strong lensing models of SMGs have a firm prediction of a few tens of strongly lensed faint SMGs per square degree (Figure 5). Since SMGs are typically very red at optical/near-infrared wavelengths they are also mainly detected at near-infrared wavelengths. As such, the near-infrared imaging from the upcoming *Euclid* survey will provide one of the best ways of finding these lensing systems. To investigate how many may be observed in the VIS and *H*-band filters of the *Euclid*-Wide survey (Laureijs et al. 2011), we perform an order-of-magnitude estimation utilising the population of SMGs in Lim et al. (2020). We restrict the sample to only Very Large Array (VLA) radio-identified SMGs, as their shorter wavelength counterparts (used to distinguish lensed SMGs) can be easily identified thanks to the higher resolution of the VLA beam compared to that of the SCUBA-2 450 μm beam. No cuts are applied to the sources, as the number density of faint SMGs vastly outnumbered that of the bright population. We interpolate their *K*-band and *L*-band AB magnitudes to *Euclid H*-band, and use the *R*-band AB magnitude for the *Euclid* VIS filter. Comparing these to the limiting AB magnitudes of *Euclid* VIS and NISP *H*-band (26.2 and 24.4 mag, respectively, Euclid Collaboration et al. 2022) suggests that only 84% of these faint SMGs will be detectable in *Euclid*-Wide VIS, but 99% will be detectable in *Euclid*-Wide NISP. For our value of 27 ± 19 strongly lensed faint SMGs per square degree within the 15000 square degree *Euclid*-Wide survey, these estimates imply a projection of 62000 ± 44000 NISP-only *Euclid*-Wide strong lenses.

Taking an alternative approach also yields a similar end result, albeit with both projections suffering from the large uncertainties incurred by our limited sample of $2 \pm \sqrt{2}$ identified lensed SMGs in the STUDIES area. For this approach, we first assume an ‘‘Eyelash’’ SED, and sample from submm source counts following Geach et al. (2013) and a redshift distribution following the model of Béthermin et al. (2012). From this, we find that at 450 μm fluxes > 5 mJy, 57% of the galaxies are detectable in NISP *H*-band above its limiting magnitude, while only 40% are detectable in VIS, with these fractions insensitive to the magnification distributions assumed. Again, based on the *Euclid*-Wide survey area and our number density, this implies a projection of 70000 ± 49000 NISP-only *Euclid*-Wide strong lenses. In summary, the implied huge number of strongly lensed *Euclid*-NISP-only galaxies is about as large as the entire strong lensing catalogue predicted for *Euclid*-VIS and famously regarded as transformative for strong lensing (e.g. Collett 2015).

Another notable large population of near-infrared-only strong gravitational lenses is the $z > 6$ lensed population. For example, Mason et al. (2015) demonstrated that the *Euclid*-Wide survey will detect ~ 100 Lyman-break galaxies at redshifts $z > 8$, nearly all of which will be strongly gravitationally lensed, but none of which will be detectable in the VIS instrument. These will appear as near-

infrared arcs around nearby $z \sim 0.5$ -2 galaxies, and will therefore be superficially very similar to the lensed SMG population, which enormously outnumber them. The degeneracy between dusty starbursts at redshifts $z < 6$ and ultra-high-redshift galaxies in *JWST* surveys has already led to downward revisions of redshift estimates (e.g. Zavala et al. 2023). This may be alleviated with the addition of photometric data from LSST, albeit at lower angular resolution for deblending the foreground lens from the background source. Promising candidates may also have dusty starburst interpretations tested by follow-up submm/mm-wave photometry (e.g. Zavala et al. 2023).

Evolved red galaxies are a further potential source of near-infrared-only lensed systems in future dark energy experiments. For example, the population of extremely red objects (EROs, e.g. $R_{\text{Johnson}} - K_{\text{Vega}} > 5$) are approximately evenly split between dusty star-forming galaxies and evolved red galaxies (e.g. Mannucci et al. 2002), with the latter typically having $H_{\text{AB}} - K_{\text{AB}} < 1$ (e.g. Bergström & Wiklind 2004). The *Euclid* $H_{\text{AB}} < 24$ limit would therefore be equivalent to a *K*-band survey at least as deep as Roche et al. (2002), implying surface densities of at least $\sim 6000 \text{ deg}^{-2}$ for each of dusty star forming galaxies and evolved red galaxies. The former is comparable to the surface density of the faintest SMGs in Figure 5 (as expected given the overlap between EROs and SMGs, e.g. Frayer et al. 2004) while the latter is clearly a comparably plentiful population of very red background sources for lensing.

5 CONCLUSIONS

It is well established that the brightest sub-millimetre galaxies (SMGs) at 450-500 μm wavelengths are almost exclusively strong gravitational lenses, with such galaxies expected to appear very red in the near-infrared. Making use of this, in this work we have searched public *JWST* NIRCcam and MIRI imaging of the COSMOS field for counterparts of SCUBA-2 450 μm sources in the STUDIES catalogue and identified a strong gravitational lensing system. Modelling of the lens indicates a mass distribution with Einstein radius $\theta_E = 0.63$ and axis ratio $q = 0.77$ that distorts the highly elliptical source into an arc with a magnification of $\mu = 3.0$. Performing SED fitting using public photometry along with the known spectroscopic redshift of the foreground lens ($z_l = 0.360$), the background source was found to be an SMG with a photometric redshift of $z_s = 3.4 \pm 0.4$.

This, along with a previously discovered SMG lens in the survey area, supports the prediction of around ten strongly lensed faint SMGs per square degree, which in turn indicates a very large population of tens of thousands of strong gravitational lenses undetectable in the optical that would be prime targets for future dark energy surveys that can image in the near-infrared, such as the *Euclid* and *Roman* wide-field surveys.

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DATA AVAILABILITY

The *Herschel* SPIRE data can be downloaded at <https://www.h-atlas.org>, while reduced and calibrated science-ready ALMA data are available on the ALMA Science Archive at <https://almascience.eso.org/asax/>. The JCMT Science Archive at the Canadian Astronomical Data Centre, <http://www.cadc-ccda.hia-ihp.nrc-cnrc.gc.ca/en/jcmt/>,

contains all raw and processed data from the JCMT’s current instrumentation. The COSMOS-Web and PRIMER calibration level 3 data, along with other publicly available *JWST* data, are available through the Space Telescope Science Institute (STScI) Mikulski Archive for Space Telescopes (MAST) Portal⁵.

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⁴ <http://www.astropy.org>

⁵ <https://mast.stsci.edu/portal/Mashup/Clients/Mast/Portal.html>

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