



Publication Year	2008
Acceptance in OA@INAF	2023-01-24T13:29:31Z
Title	The high activity of 3C 454.3 in autumn 2007: Monitoring by the WEBT during the AGILE detection
Authors	RAITERI, Claudia Maria; VILLATA, Massimo; Chen, W.P.; Hsiao, W.-S.; Kurtanidze, O.M.; et al.
DOI	10.1051/0004-6361:200809995
Handle	http://hdl.handle.net/20.500.12386/33038
Journal	ASTRONOMY & ASTROPHYSICS
Number	485

LETTER TO THE EDITOR

The high activity of 3C 454.3 in autumn 2007 Monitoring by the WEBT during the AGILE detection^{*}

C. M. Raiteri¹, M. Villata¹, W. P. Chen², W.-S. Hsiao², O. M. Kurtanidze³, K. Nilsson⁴, V. M. Larionov^{5,6}, M. A. Gurwell⁷, I. Agudo⁸, H. D. Aller⁹, M. F. Aller⁹, E. Angelakis¹⁰, A. A. Arkharov⁶, U. Bach¹⁰, M. Böttcher¹¹, C. S. Buemi¹², P. Calcidese¹³, P. Charlot¹⁴, F. D'Ammando¹⁵, I. Donnarumma¹⁵, E. Forné¹⁶, A. Frasca¹², L. Fuhrmann¹⁰, J. L. Gómez⁸, V. A. Hagen-Thorn⁵, S. G. Jorstad¹⁷, G. N. Kimeridze³, T. P. Krichbaum¹⁰, A. Lähteenmäki¹⁸, L. Lanteri¹, G. Latev¹⁹, J.-F. Le Campion¹⁴, C.-U. Lee²⁰, P. Leto²¹, H.-C. Lin², N. Marchili¹⁰, E. Marilli¹², A. P. Marscher¹⁷, R. Nesci²², E. Nieppola¹⁸, M. G. Nikolashvili³, J. Ohlert²³, E. Ovcharov¹⁹, D. Principe¹¹, T. Pursimo²⁴, B. Ragozzine¹¹, A. C. Sadun²⁵, L. A. Sigua³, R. L. Smart¹, A. Strigachev²⁶, L. O. Takalo⁴, M. Tavani¹⁵, C. Thum²⁷, M. Tornikoski¹⁸, C. Trigilio¹², K. Uckert¹¹, G. Umama¹², A. Valcheva²⁶, S. Vercellone²⁸, A. Volvach²⁹, and H. Wiesemeyer³⁰

(Affiliations can be found after the references)

Received 17 April 2008 / Accepted 9 May 2008

ABSTRACT

Context. The quasar-type blazar 3C 454.3 underwent a phase of high activity in summer and autumn 2007 that was intensively monitored in the radio-to-optical bands by the Whole Earth Blazar Telescope (WEBT). The γ -ray satellite Astro-rivelatore Gamma a Immagini LEggero (AGILE) detected this source first in late July, and then in November–December 2007.

Aims. We present the multifrequency data collected by the WEBT and collaborators during the second AGILE observing period, complemented by contemporaneous data from the UltraViolet and Optical Telescope (UVOT) onboard the Swift satellite. The aim is to trace in detail the behaviour of the synchrotron emission from the blazar jet, and to investigate the contribution from the thermal emission component.

Methods. Optical data from about twenty telescopes have been homogeneously calibrated and carefully assembled to construct an *R*-band light curve containing about 1340 data points over 42 days. This extremely well-sampled optical light curve allows us to follow the dramatic flux variability of the source in detail. In addition, we show radio-to-UV spectral energy distributions (SEDs) at different epochs, which represent different brightness levels.

Results. In the considered period, the source varied by 2.6 mag over two weeks in the *R* band. Many episodes of fast (i.e. intranight) variability were observed, most notably on December 12, when a flux increase of about 1.1 mag in 1.5 h was detected, followed by a steep decrease of about 1.2 mag in 1 h. The contribution by the thermal component is difficult to assess, due to the uncertainties in the Galactic, and possibly also intrinsic, extinction in the *UV* band. However, polynomial fitting of radio-to-UV SEDs reveals an increasing spectral bending towards fainter states, suggesting a UV excess likely due to the thermal emission from the accretion disc.

Conclusions. Once the AGILE data are completely analysed, the low-frequency observations presented in this letter will offer a powerful tool to investigate optical- γ flux correlations, i.e. the relationship between the synchrotron and inverse-Compton emission components.

Key words. galaxies: active – galaxies: quasars: general – galaxies: quasars: individual: 3C 454.3

1. Introduction

The flat-spectrum radio quasar 3C 454.3 caught the attention of astronomers in spring 2005, when it was observed in an exceptional brightness state from the near-IR to the X-ray frequency range. The Whole Earth Blazar Telescope (WEBT)¹ collaboration organised a multiwavelength campaign to study the event, whose results were published by Villata et al. (2006). The WEBT monitoring effort continued after the optical outburst, and followed the subsequent radio activity (Villata et al. 2007), and then the faint state in the 2006–2007 observing season (Raiteri et al. 2007). In this last period, the low contribution of the synchrotron emission from the jet allowed the authors to recognise both the

little blue bump, due to line emission from the broad line region, and the big blue bump, due to thermal emission from the accretion disc.

Renewed optical activity was observed at the beginning of the next observing season, in May 2007, which prompted the WEBT collaboration to continue with the monitoring. Indeed, a strong optical outburst was observed in July–August. This triggered observations by the Astro-rivelatore Gamma a Immagini LEggero (AGILE) satellite, which detected the source in its brightest γ -ray state ever observed (Vercellone et al. 2008). A further significant detection by AGILE was announced in November (Chen et al. 2007; Pucella et al. 2007), which in turn led to intensified monitoring by the WEBT.

The importance of collecting data simultaneously at low and high energies comes from the fact that the X-to- γ -ray radiation from blazars is commonly thought to be produced in the plasma jet through an inverse-Compton process off the same relativistic electrons that generate the radio-to-optical synchrotron

^{*} The radio-to-optical data presented in this paper are stored in the WEBT archive; for questions regarding their availability, please contact the WEBT President Massimo Villata.

¹ <http://www.oato.inaf.it/blazars/webt/>
see e.g. Villata et al. (2004); Böttcher et al. (2007); Raiteri et al. (2008).

emission. The seed photons for the inverse-Compton scattering may be either the same synchrotron photons (synchrotron-self-Compton or SSC models) or photons coming from outside the jet (external-Compton or EC models), or a combination of both. In particular, the γ -ray radiation in the GeV regime should be directly related to the optical one. Hence, the comparison between the γ -ray AGILE data and the optical data, especially during periods of high variability, can constitute an important contribution to our knowledge of blazar emission mechanisms. An example of synergy between AGILE and WEBT has recently been shown on the blazar S5 0716+71 (Villata et al. 2008).

In this letter we deal with the multiwavelength data taken in the core period (November 4–December 16, 2007) of the last WEBT campaign on 3C 454.3. This core period corresponds to the AGILE second detection, and includes a few observations by the Swift satellite. We will present a detailed analysis of all multiwavelength data acquired during the last WEBT campaign (May 2007–February 2008) in a forthcoming paper (Raiteri et al., in prep.).

2. Optical-to-radio observations by the WEBT

Optical and near-IR data were acquired in Johnson-Cousins *UBVRI* and *JHK* bands by about twenty telescopes all around the world. Magnitudes were calibrated following Angione (1971) for the *U* filter, Raiteri et al. (1998) for the *BVR* filters, and González-Pérez et al. (2001) for the *I* filter. The *JHK* data were calibrated according to a photometric sequence by Larionov (private commun.). The light curve construction was performed by assembling all datasets for each band and then simultaneously inspecting all the resulting optical-NIR light curves day by day. When an intranight dataset contained noisy data, they were binned (typically in 5 min bins); redundant data points with too large errors (≥ 0.1 mag) and unreliable outliers were discarded.

The *R*-band light curve obtained during the core period of the WEBT campaign is shown in Fig. 1. It was constructed with the data provided by the following observatories (ordered by decreasing contribution): Lulin, Abastumani, Roque de los Muchachos (KVA and NOT), St. Petersburg, Mt. Lemmon, Valle d’Aosta, Torino, Kitt Peak (MDM), L’Ampolla, Vallinfreda, Michael Adrian, Rozhen (50/70 cm and 200 cm telescopes), New Mexico Skies, and Calar Alto². The light curve contains 1337 data points, which show a total variation of 2.58 mag, occurring during the last 15.18 days. The Bordeaux and Catania observatories also contributed to the optical light curves, but not in the *R* band.

The source was very active in this period, and strong variations were also observed on intranight time scales. Noticeable episodes of fast variability are displayed in Fig. 2. Figure 2a shows a well-sampled brightening of 0.33 mag in 2.3 h, which was followed by an apparently slower flux decrease. The first dimming phase in Fig. 2b implies a variation of 0.38 mag in 9.0 h. A very spectacular variability episode is shown in Fig. 2c, where the total flux increase is 1.35 mag in 28.5 h, including an almost monotonic brightening of 0.83 mag in 4.6 h³, from JD = 2 454 435.9357 to 2 454 436.1284, corresponding to

² Calar Alto data were taken as part of the MAPCAT (Monitoring AGN with Polarimetry at the Calar Alto Telescopes) project.

³ With very steep slopes of about 0.006 mag per minute, i.e. 3 times steeper than the fastest variations observed in the strong IDV blazar 0716+714 (see Villata et al. 2008, and references therein).

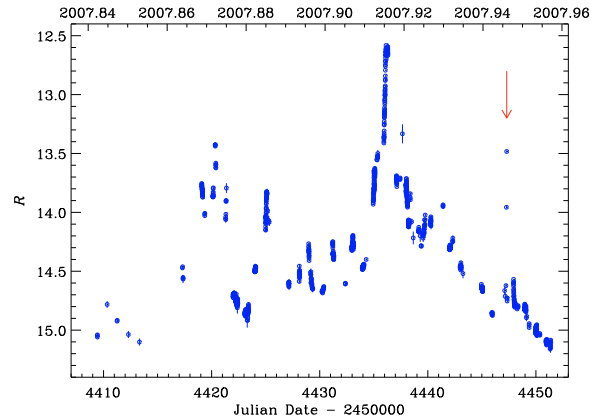


Fig. 1. *R*-band light curve of 3C 454.3 in November 4–December 16, 2007. The episode of extremely fast variability indicated by the red arrow is shown in more detail in Fig. 3.

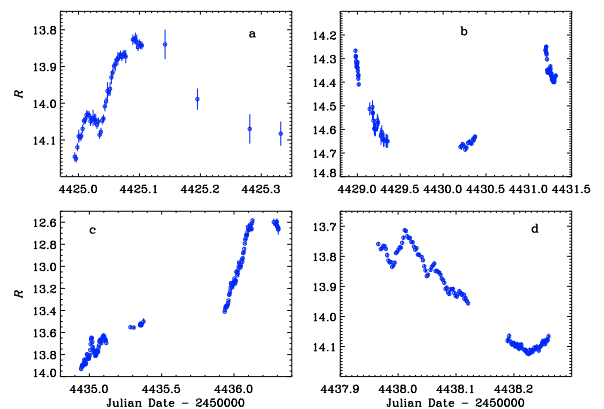


Fig. 2. Examples of very fast variability of 3C 454.3 extracted from the light curve in Fig. 1.

the maximum peak of the light curve in Fig. 1. The flux variations displayed in Fig. 2d involve a 0.41 mag dimming in 5.1 h.

There was an episode of extraordinarily fast variability detected on December 12, 2007, which is shown in Fig. 3. Observations started in St. Petersburg (*BVRI*) and then continued in Valle d’Aosta (*BVRI*) and Bordeaux (*V*). For all these datasets the source was relatively faint. But soon after, eight *BVRI* frames taken at the Torino Observatory revealed a sudden brightening, which was over when the Roque (KVA) *R*-band data were acquired. This rapid flare implies at least a 1.1 mag brightening in 1.5 h, and a brightness decrease of about 1.2 mag in 1 h in the *R* band⁴, thus representing one of the most dramatic flux changes ever detected in a blazar. For this reason the Torino frames were carefully analysed to check for any possible instrumental problem and artifact, finding none. Hence, we believe that the observed fast variation represents a real phenomenon.

At the beginning of the period considered in this letter, a few near-IR data were taken at Campo Imperatore and with the Nordic Optical Telescope (NOT).

Radio data at both mm and cm wavelengths were collected as already calibrated flux densities. In the mm regime, data were taken at 850 μ m and 1 mm with the Submillimeter Array (SMA) on Mauna Kea⁵, and at 1 mm, 142 GHz, and 86 GHz with the

⁴ More precisely, 1.25 mag in 64 min, i.e. a mean slope of 0.020 mag per minute (cf. Footnote 3).

⁵ These data were obtained as part of the normal monitoring program initiated by the SMA (Gurwell et al. 2007).

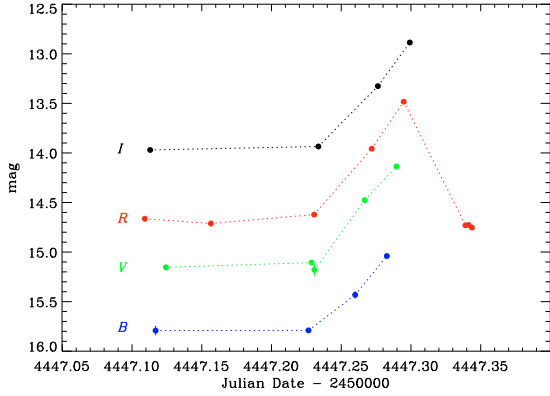


Fig. 3. The extraordinary episode of fast variability observed on December 12, 2007.

IRAM 30 m telescope on Pico Veleta⁶. At cm wavelengths, the radio flux densities were measured at the Metsähovi (37 GHz), UMRAO (14.5, 8.0, and 4.8 GHz), Effelsberg (42, 23, 15, 8, 5, 3, and 1 GHz), Crimean (37 GHz), Medicina (22, 8, and 5 GHz), and Noto (43 GHz) observatories. We also collected data from the VLA/VLBA Polarization Calibration Database⁷.

3. UV data from Swift and spectral energy distributions

In the period considered in this letter, observations by the UVOT instrument (Roming et al. 2005) onboard Swift were performed on November 15, 20, and 22, and on December 13 and 15. However, only in the first and in the last two epochs all the optical (*UBV*) and UV (*W2*, *M2*, *W1*) filters were used, while in the other epochs data in the *UV* bands only were acquired. The UVOT data were processed with the `uvotmaghist` task of the HEASOFT 6.3 package. Following the recommendations contained in the release notes, the source counts were extracted from a circular region with a 5 arcsec radius, in agreement with the standard photometric aperture defined in the calibration files (CALDB updated as July 2007). To avoid contamination by Star 1 (see Raiteri et al. 1998), the background was extracted from source-free circular regions in the source surroundings. We attributed a 0.1 mag uncertainty to all UVOT data, to take account of both systematic and statistical errors.

We constructed broad-band SEDs with contemporaneous data from the radio bands to the UV frequencies. We selected epochs in which data at various frequencies were available, to better trace the synchrotron component. In Fig. 4 we show the most significant SEDs at different brightness levels. Before being converted into flux densities, NIR-to-UV magnitudes were de-reddened according to the Cardelli et al. (1989) laws, assuming an extinction in the *B* band of 0.462 mag, and $R_V = A(V)/E(B - V) = 3.1$, as in the diffuse interstellar medium. One delicate point is the amount of extinction in the *UV* bands. The UVOT effective wavelengths for the *UVW1*, *UVM2*, and *UVW2* filters are 2634, 2231, and 2030 Å, respectively (Poole et al. 2008). We derive for these bands an extinction of 0.73, 1.07, and 1.02 mag, respectively, resulting in the curved shape

⁶ The IRAM data are part of a standard AGN monitoring program of total flux and polarization, a complementary program to monitor additional sources selected for observations with GLAST (Fuhrmann et al., in prep.), and a dedicated polarimetry program to monitor the mm flare of 3C 454.3 in 2007 (Agudo et al., in prep.).

⁷ <http://www.vla.nrao.edu/astro/calib/polar/>

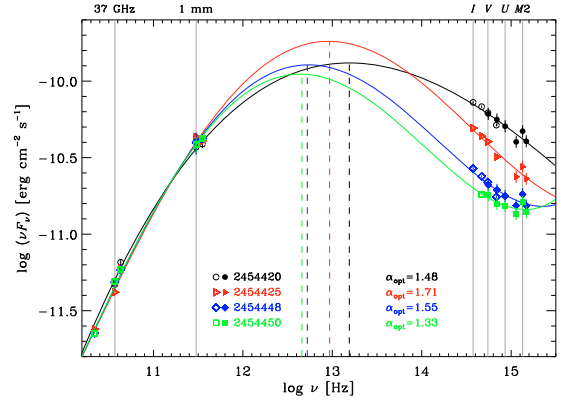


Fig. 4. Spectral energy distributions (SEDs) of 3C 454.3 at different brightness levels in November–December 2007. Contemporaneous data are plotted with the same symbol and colour; filled symbols refer to Swift-UVOT data, while empty symbols represent ground-based observations. The curves show polynomial fits to the radio-to-UV data at the various epochs; dashed vertical lines correspond to the fit maxima. The values of the optical spectral indices are also indicated.

of the SEDs in Fig. 4. We notice that the amount of Galactic extinction in the UV is strongly affected by the 2175 Å bump due to dust absorption. Even slight variations in the depth and/or shape of this feature would have noticeable consequences on the SEDs in the UV, particularly on the position of the *UVM2* points, since the effective wavelength of this filter is close to 2175 Å. Moreover, fitting the 3C 454.3 X-ray spectra requires extra absorption (see e.g. Raiteri et al. 2007, and references therein). Hence, there may be intrinsic absorption also in the optical-UV frequency range (see e.g. Gaskell et al. 2004), even if it may have different properties from the Galactic extinction (Maiolino et al. 2001).

Consequently, the shape of the spectrum in the *UV* band has to be taken with caution. However, if we perform a polynomial fit of the SEDs from the radio (starting from 5 GHz, outside the plot in Fig. 4) to the UV data, we can see that the optical-UV SED becomes more and more concave as the flux decreases, suggesting the presence of a rather stable UV excess. Indeed, a UV excess was found by Raiteri et al. (2007) when analysing XMM-Newton data taken in July and December 2006, which was interpreted as the contribution from the accretion disc. Our SEDs would confirm the picture according to which the thermal contribution from the disc is overwhelmed by the beamed synchrotron emission when the source is bright, and it can be recognised only when the non-thermal flux is low.

We finally notice the remarkable change in the optical spectral index α_{opt} ($F_\nu \propto \nu^{-\alpha}$), from 1.33 (JD = 2454450) to 1.71 (JD = 2454425), and that the polynomial fits suggest that the peak of the SED may shift towards higher frequencies with increasing optical flux (see Fig. 4).

4. Discussion and conclusions

In this letter we have presented the radio-to-optical data on 3C 454.3 acquired by the WEBT and collaborators, complemented by data taken with the UVOT instrument onboard Swift, during the core period of the last WEBT campaign, i.e. from November 4 to December 16, 2007. The source showed strong activity in the optical band, with a maximum variation of about 2.6 mag over two weeks. Many episodes of fast variability were observed, involving flux changes of several tenths of mag in a few hours. The most striking event was detected on

December 12, when a flux increase of about 1.1 mag in 1.5 h was followed by a fall of about 1.2 mag in 1 h. This is one of the fastest variations ever observed in blazars, even if the inferred brightness temperature is still well below the limit of $\sim 10^{12}$ K for an inverse-Compton catastrophe. Indeed, by assuming $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$, the 1.2 mag dimming in 1 h in the R band yields $T_b \sim 5 \times 10^9 \text{ K}$ (e.g. Wagner et al. 1996)⁸. We think that this very fast flare is most likely a real variation in the jet emission. Alternatively, it may be due to a dramatic event in the host galaxy, like a γ -ray burst (GRB). However, most GRBs are associated with the death of young massive stars, which are more common in star-forming galaxies (Savaglio et al. 2008), while blazars are usually hosted in ellipticals. A microlensing effect, e.g. by a MACHO in our Galaxy, may be possible, but it seems unlikely; indeed, these events are observed with much longer time scales (e.g. Wood & Mao 2005). Unfortunately, the poor sampling prevents us from drawing any conclusions on the shape and chromatism of the event.

The construction of broad-band SEDs with contemporaneous data reveals noticeable spectral changes in the optical-UV frequency range. Although the details of the spectral shape must be regarded with caution because of the uncertainties in the modelling of absorption (Galactic, but possibly also intrinsic), especially in the UV band, we notice that the optical-UV SED becomes more concave as the flux decreases. This is consistent with the finding of a UV excess in faint states (Raiteri et al. 2007), likely due to thermal radiation from the accretion disc. Indeed, this contribution to the blazar emission becomes more visible when the beamed synchrotron radiation from the jet is fainter, which stresses the importance of studying blazars in low-activity states as well as during outbursts.

The time interval considered in this letter corresponds to the second period of detection in γ -rays by the AGILE satellite, which follows the observation of an unprecedented bright γ -ray state in July 2007 (Vercellone et al. 2008). The AGILE data in November–December 2007 show high variability on time scales from days to weeks, with recurrent flaring activity (Vercellone et al., in prep.; Donnarumma et al., in prep.). The comparison between the low-energy data presented in this letter and the high-energy data by AGILE will contribute to shedding light on the relationship between the synchrotron and inverse-Compton emission components from the blazar jet.

Acknowledgements. We acknowledge the use of public data from the Swift data archive. This research has made use of data obtained through the High Energy Astrophysics Science Archive Research Center Online Service, provided by the NASA/Goddard Space Flight Center. This work is partly based on observations made with the Nordic Optical Telescope, operated on the island of La Palma jointly by Denmark, Finland, Iceland, Norway, and Sweden, in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias, and on observations collected at the German-Spanish Calar Alto Observatory, jointly operated by the MPIA and the IAA-CSIC. AZT-24 observations are made within an agreement between Pulkovo, Rome and Teramo observatories. We thank the staff at the IRAM 30-m telescope for their help with the observations, and in particular H. Ungerechts for allowing us to include data from the general IRAM 30-m AGN monitoring program. We are also grateful to the IRAM Director for providing discretionary observing time at the 30-m telescope. This paper is based on observations carried out at the IRAM 30-m telescope. IRAM is supported by INSU/CNRS (France), MPG (Germany) and IGN (Spain). This research has made use of data from the University of Michigan Radio Astronomy Observatory, which is supported by the National Science Foundation and NASA and by funds from the University of Michigan. The Metsähovi team acknowledges the support from the Academy of Finland. The Torino team acknowledges financial support by the Italian Space Agency through contract ASI-INAF I/088/06/0 for the Study of High-Energy Astrophysics.

Acquisition of the MAPCAT data is supported in part by the Spanish Ministerio de Educacion y Ciencia through grant AYA2007-67626-C03-03.

References

- Angione, R. J. 1971, *AJ*, 76, 412
 Böttcher, M., Basu, S., Joshi, M., et al. 2007, *ApJ*, 670, 968
 Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245
 Chen, A., Vercellone, S., Giuliani, A., et al. 2007, *The Astronomer's Telegram*, 1278, 1
 Gaskell, C. M., Goosmann, R. W., Antonucci, R. R. J., & Whysong, D. H. 2004, *ApJ*, 616, 147
 González-Pérez, J. N., Kidger, M. R., & Martín-Luis, F. 2001, *AJ*, 122, 2055
 Gurwell, M. A., Peck, A. B., Hostler, S. R., Darrah, M. R., & Katz, C. A. 2007, in *From Z-Machines to ALMA: (Sub)Millimeter Spectroscopy of Galaxies*, ed. A. J. Baker, J. Glenn, A. I. Harris, J. G. Mangum, & M. S. Yun, *ASP Conf. Ser.*, 375, 234
 Maiolino, R., Marconi, A., Salvati, M., et al. 2001, *A&A*, 365, 28
 Poole, T. S., Breeveld, A. A., Page, M. J., et al. 2008, *MNRAS*, 383, 627
 Pucella, G., Tavani, M., D'Ammando, F., et al. 2007, *The Astronomer's Telegram*, 1300, 1
 Raiteri, C. M., Villata, M., Lanteri, L., Cavallone, M., & Sobrito, G. 1998, *A&AS*, 130, 495
 Raiteri, C. M., Villata, M., Larionov, V. M., et al. 2007, *A&A*, 473, 819
 Raiteri, C. M., Villata, M., Larionov, V. M., et al. 2008, *A&A*, 480, 339
 Roming, P. W. A., Kennedy, T. E., Mason, K. O., et al. 2005, *Space Sci. Rev.*, 120, 95
 Savaglio, S., Glazebrook, K., & Le Borgne, D. 2008, [[arXiv:0803.2718](https://arxiv.org/abs/0803.2718)]
 Vercellone, S., Chen, A. W., Giuliani, A., et al. 2008, *ApJ*, 676, L13
 Villata, M., Raiteri, C. M., Aller, H. D., et al. 2004, *A&A*, 424, 497
 Villata, M., Raiteri, C. M., Balonek, T. J., et al. 2006, *A&A*, 453, 817
 Villata, M., Raiteri, C. M., Aller, M. F., et al. 2007, *A&A*, 464, L5
 Villata, M., Raiteri, C. M., Larionov, V. M., et al. 2008, *A&A*, 481, L79
 Wagner, S. J., Witzel, A., Heidt, J., et al. 1996, *AJ*, 111, 2187
 Wood, A., & Mao, S. 2005, *MNRAS*, 362, 945

¹ INAF, Osservatorio Astronomico di Torino, Italy

e-mail: [raiteri;villata]@oato.inaf.it

² Institute of Astronomy, National Central University, Taiwan

³ Abastumani Astrophysical Observatory, Georgia

⁴ Tuorla Observatory, Univ. of Turku, Piikkiö, Finland

⁵ Astron. Inst., St.-Petersburg State Univ., Russia

⁶ Pulkovo Observatory, St. Petersburg, Russia

⁷ Harvard-Smithsonian Center for Astroph., Cambridge, MA, USA

⁸ Instituto de Astrofísica de Andalucía (CSIC), Granada, Spain

⁹ Department of Astronomy, University of Michigan, MI, USA

¹⁰ Max-Planck-Institut für Radioastronomie, Bonn, Germany

¹¹ Department of Physics and Astronomy, Ohio Univ., OH, USA

¹² INAF, Osservatorio Astrofisico di Catania, Italy

¹³ Osservatorio Astronomico della Regione Autonoma Valle d'Aosta, Italy

¹⁴ Lab. d'Astrophysique, Université Bordeaux 1, CNRS, Floirac, France

¹⁵ INAF, IASF-Roma, Italy

¹⁶ Agrupació Astronòmica de Sabadell, Spain

¹⁷ Institute for Astrophysical Research, Boston University, MA, USA

¹⁸ Metsähovi Radio Obs., Helsinki Univ. of Technology TKK, Finland

¹⁹ Sofia University, Bulgaria

²⁰ Korea Astronomy and Space Science Institute, South Korea

²¹ INAF, Istituto di Radioastronomia, Sezione di Noto, Italy

²² Dept. of Phys. "La Sapienza" Univ, Roma, Italy

²³ Michael Adrian Observatory, Trebur, Germany

²⁴ Nordic Optical Telescope, Santa Cruz de La Palma, Spain

²⁵ Dept. of Phys., Univ. of Colorado, Denver, USA

²⁶ Inst. of Astronomy, Bulgarian Academy of Sciences, Sofia, Bulgaria

²⁷ Institut de Radio Astronomie Millimétrique, Grenoble, France

²⁸ INAF, IASF-Milano, Italy

²⁹ Radio Astronomy Lab. of Crimean Astrophysical Observatory, Ukraine

³⁰ Instituto de Radioastronomía Millimétrica, Granada, Spain

⁸ For comparison, the 0.83 mag brightening in 4.6 h observed on December 1 (JD $\sim 2\,454\,436.0$, see Fig. 2) implies $T_b \sim 4 \times 10^8 \text{ K}$.