


Barium lines in high-quality spectra of two metal-poor giants in the Galactic halo[★]

G. Cescutti^{1,2,3}, C. Morossi¹, M. Franchini¹, P. Di Marcantonio¹, C. Chiappini⁴, M. Steffen⁴, M. Valentini⁴, P. François^{5,6}, N. Christlieb⁷, C. Cortés^{8,9}, C. Kobayashi¹⁰, and E. Depagne¹¹

¹ INAF, Osservatorio Astronomico di Trieste, Via Tiepolo 11, 34143 Trieste, Italy
e-mail: gabriele.cescutti@inaf.it

² IFPU, Institute for the Fundamental Physics of the Universe, Via Beirut, 2, 34151 Grignano, Trieste, Italy

³ INFN, Sezione di Trieste, Via A. Valerio 2, 34127 Trieste, Italy

⁴ Leibniz-Institut für Astrophysik Potsdam (AIP), An der Sternwarte 16, 14482, Potsdam, Germany

⁵ GEPI – Observatoire de Paris, 64 avenue de l’Observatoire, 75014, Paris, France

⁶ Université de Picardie Jules Verne, 33 rue St-Leu, 80080 Amiens, France

⁷ Zentrum für Astronomie der Universität Heidelberg, Landessternwarte, Königstuhl 12, 69117 Heidelberg, Germany

⁸ Departamento de Física, Facultad de Ciencias Básicas, Universidad Metropolitana de la Educación, Av. José Pedro Alessandri 774, 7760197 Nuñoa, Santiago, Chile

⁹ Millennium Institute of Astrophysics (MAS), Santiago, Chile

¹⁰ Centre for Astrophysics Research, Department of Physics, Astronomy and Mathematics, University of Hertfordshire, College Lane, Hatfield AL10 9AB, UK

¹¹ South African Astronomical Observatory (SAAO), Observatory Road Observatory Cape Town, WC 7925, South Africa

Received 20 May 2021 / Accepted 11 August 2021

ABSTRACT

Context. Theoretical results showed the possibility that neutron capture elements were produced in the early Universe by two different sources: a frequent *s*-process source hosted by rotating massive stars, and a rare *r*-process source hosted most likely by neutron star mergers. The two sources produce barium with different isotopic compositions.

Aims. We aim to investigate the lines of barium in two halo stars, HD 6268 and HD 4306. The spectra present an exquisite quality, both in terms of resolution ($R > 100\,000$) and signal-to-noise (~ 400). Due to hyperfine splitting (hfs) effects, barium lines are expected to show slightly different profiles depending on the barium isotopic fraction.

Methods. We applied a standard local thermodynamic equilibrium synthesis of the barium lines. We compared the synthetic results assuming an *s*-process isotopic pattern or an *r*-process isotopic pattern for the two barium lines for each star that exhibited hfs. We also applied a methodology, less dependent on the accuracy of the theoretical Ba hfs structure, that transforms the lines of HD 4306 into those we would observe if its atmospheric parameter values (i.e. T_{eff} , $\log g$, micro- and macro-turbulence, $V \sin i$, and Ba abundance) were the same as those of HD 6268.

Results. With both methods, our results show that the barium lines with hfs effects of HD 4306 are in agreement with an *s*-process composition and the lines in HD 6268 have a different profile, which is most likely linked to the presence of an *r*-process isotopic pattern.

Conclusions. Two lines of barium of HD 6268 and HD 4306 seem to confirm the theoretical expectation that both *r*-process events and also *s*-process contribution by rotating massive stars have polluted the ancient halo of our Galaxy.

Key words. Galaxy: evolution – Galaxy: halo – stars: abundances – stars: rotation – nuclear reactions, nucleosynthesis, abundances

1. Introduction

The past decade has seen an explosion of elemental abundance determinations for the most metal-poor stars in the Galaxy. These new data have shown many surprising results, including the fact that light neutron capture elements (such as Sr, Y, and Zr) are more enriched in some extremely metal-poor (EMP) stars compared to the expectations of a pure *r*-process pattern (Thielemann et al. 2011), and they show an increasing scatter towards low metallicities, down to $[\text{Fe}/\text{H}]^1 \sim -4$ (Honda et al.

2004; Barklem et al. 2005; François et al. 2007; Spite et al. 2018).

Astronomers then invoked the need for a new process (e.g. Travaglio et al. 2004) able to provide the additional contribution to light neutron capture elements. The astrophysical conditions that would create this additional contribution are still unknown, and both an *s*-process-like or an *r*-process-like mechanism was found to reproduce the abundance pattern between Sr and Ag observed in many of the most extreme metal poor stars (Montes et al. 2007). The *s*-process production in massive stars (weak *s*-process) was shown to be inefficient at low metallicity (Raiteri et al. 1992), so most authors have focused on the existence of a possible second *r*-process (also called the weak *r*-process; see Arcones & Montes 2011).

However, the recent nucleosynthesis computations by Frischknecht et al. (2016) showing that rotating massive stars

[★] Based on observations made with the UVES at the ESO Very Large Telescope, Paranal Observatory, Chile (ID 098.B-0094(A); P.I. G. Cescutti).

¹ We assumed the common approximation in which iron is a proxy for the entire metal content, and we adopted the notation $[A/B] \equiv \log_{10}(N_A/N_B) - \log_{10}(N_A/N_B)_\odot$ and $\epsilon(A) \equiv \log_{10}(N_A/N_H)$ for elements A and B.

can support the *s*-process, and hence pollute the early Universe with these elements on a timescale comparable to that of the *r*-process, have brought a new twist to the interpretation of the neutron capture element abundances of stars at very low metallicity. Interestingly, the existence of rotating massive stars in the early Universe has already been suggested as the only solution for at least four ‘anomalies’ observed in other chemical elements in very metal-poor stars ($[\text{Fe}/\text{H}] < -2.5$); namely, the rise of N/O and C/O towards lower metallicities (Chiappini et al. 2006); the low $^{12}\text{C}/^{13}\text{C}$ (Chiappini et al. 2008) and the primary-like behaviour of Be and B (Prantzos 2012). Recently, we claimed to have found a fifth signature of the impact of rotating massive stars, namely an early enrichment of the Universe in *s*-process elements (Chiappini et al. 2011; Cescutti et al. 2013). More recently, the impact of rotation in the nucleosynthesis of neutron capture elements was also confirmed by the studies of Limongi & Chieffi (2018) and Choplin et al. (2018). Moreover, chemical evolution models have again confirmed that adopting these yields of rotating massive stars can explain the increase in $[\text{Sr}/\text{Ba}]$ towards low metallicity (Prantzos et al. 2018; Rizzuti et al. 2019, 2021). A similar increase was obtained by Kobayashi et al. (2020) thanks to the strontium production by electron capture SNe, a potential source of weak *r*-process; in this case, the barium is formed only by an *r*-process.

In this framework, the oldest halo stars can form from an ISM enriched by an *r*-process source or an *s*-process source (or both). A clear prediction of stochastic models with rotating massive stars is that EMP stars with high $[\text{Sr}/\text{Ba}]$ ratio should be almost entirely enriched by the *s*-process (see Fig. 2 in Cescutti & Chiappini 2014). On the other hand, an *r*-process enrichment is expected for stars that show a lower $[\text{Sr}/\text{Ba}]$ ratio, closer to the *r*-process solar residual² ratio (Arlandini et al. 1999). Following this interpretation, the abundances of the isotopes of the elements produced are expected to present different patterns among these two classes of stars. Usually, atomic lines of heavy elements do not allow us to distinguish the isotopic composition of the element. However, Ba is an attractive heavy element in this respect. In fact, it presents a hyperfine splitting (hfs) of its 4554 Å line and its 4934 Å line from the singly ionised stage, which is large enough to be detected (Rutten 1978) and hence offers the possibility of measuring the fraction of odd isotopes (hereafter, f_{odd} ³) via resolved asymmetric lines. According to nucleosynthesis calculations (Arlandini et al. 1999), one expects an $f_{\text{odd}} = 0.11 \pm 0.01$ in the case of a pure *s*-process, and an $f_{\text{odd}} = 0.46 \pm 0.06$ in the case of the *r*-process solar residual. So, in the case of an *s*-process enrichment, even isotopes and no hfs effects are almost exclusively expected; on the other hand, in the case of an *r*-process the impact of the profile of the barium lines is relatively strong and driven by the large fraction of odd isotopes prone to hfs effects. Kobayashi et al. (2020) found a higher $f_{\text{odd}} = 0.3$ for their chemical evolution model with rotating massive stars at $[\text{Fe}/\text{H}] = -2.7$; this is due to the model, which is in fact a homogeneous one that mixes the contribution of the *r*-process events and rotating massive stars. The same work, based on recent theoretical computations of nucleosynthesis, found a slightly higher fraction of odd isotopes ($f_{\text{odd}} = 0.63$) for the pure *r*-process.

The measurement of the Ba-isotopic ratio, although feasible, is very challenging. Magain (1995) measured the isotopic ratio of a very bright halo star, HD 140283, for the first time. He

found an *s*-process signature with a high $[\text{Sr}/\text{Ba}] = 0.9$, which was in agreement with the theoretical expectations of the chemical models of Cescutti & Chiappini (2014). However, this result on HD 140283 has been challenged and still needs to be confirmed (see Lambert & Allende Prieto 2002; Collet et al. 2009; Gallagher et al. 2010). Gallagher et al. (2012) attempted to measure isotopic ratios in five *s*-process-dominated metal-poor stars, obtaining isotopic fractions more compatible with an *s*-process-dominated composition. Indeed, one of the biggest difficulties in this case is properly accounting for 3D effects on the line formation. Gallagher et al. (2015) adopted a 3D modelling to obtain an *r*-process isotopic composition for HD 140283, contrary to their previous results in 1D and also differently to the similar attempt in 3D carried out by Collet et al. (2009). In the same year, Jablonka et al. (2015) also managed to determine this isotopic fraction in stars belonging to the dwarf spheroidal Sculptor.

In this work, we investigated two spectra of extremely high quality (signal-to-noise, $S/N > 400$ and $R \sim 100\,000$) for two stars, HD 4306 and HD 6268. The spectra were obtained homogeneously with the same instrument, which was set up and during the same night. The final goal is to compare the barium lines of these two stars. For this purpose, the stars were selected to have parameters as similar as possible: $[\text{Fe}/\text{H}] \sim -2.5$ dex, $\log g \sim 1$, and $T_{\text{eff}} \sim 4500$ K, but with a different ratio of $[\text{Sr}/\text{Ba}]$; so, we expect HD 4306 to show barium with an isotopic pattern typical of an *s*-process, and HD 6268 to show that typical of an *r*-process enrichment. The parameters were not as close as ideally desired, but the constraints of bright halo stars with $[\text{Fe}/\text{H}] < -2.5$ resulted in the choice of these two candidates. The paper is structured as follows: in Sect. 2, we describe the observations and the reductions of the stellar spectra; in Sect. 3, we provide new atmospheric parameters, with a comparison to the previous literature results; in Sect. 4, we present the chemical abundances obtained from equivalent widths; in Sect. 5, we show the comparison of the synthetic spectra with *r*- and *s*-process isotopic compositions to the observational ones; in Sect. 6, a new empirical method is described to directly compare the two observational spectra; in Sect. 7, we draw our conclusions.

2. Observations and data reduction

The observations were performed using the UVES high-resolution spectrograph (Dekker et al. 2000), mounted at the UT2 Kueyen Telescope at the ESO Paranal Observatory (Chile), in slit mode. We adopted the standard R530 set-up (red arm only, cross disperser 3, centred at 5200 covering the 4140–6210 Å wavelength interval with a resolving power of $R \sim 100\,000$).

The spectra were acquired on the night of 18th October 2016, with three exposures of 2400 s (40 min) for HD 4306 and two exposures of 1600 s (26 min and 40 s) for HD 6268. Each of the resulting spectra have approximately an average S/N of 250 for HD 4306 and 350 for HD 6268. Details of the observations are summarised in Table 1.

The spectra were reduced using the standard UVES pipeline with the following exceptions:

- 23 high-quality flat fields were purposely obtained and used to build the MASTERFLAT;
- wavelength calibration was achieved by using ThAr lamps obtained immediately after the stellar observations;
- flat-fielding is obtained in pixel-pixel space before extraction (`reduce.ffmethod = pixel`) and not using the default option ‘extract’;

² The residual abundances of chemical species of the solar abundances; after that, theoretical *s*-process abundances have been removed.

³ $f_{\text{odd}} = [N(^{135}\text{Ba}) + N(^{137}\text{Ba})]/N(\text{Ba})$.

Table 1. Observing log.

Object	RA (J2000.0)	Dec (J2000.0)	B (mag)	G (mag)	Exp. time (s)	$\langle S/N \rangle$	Obs. date (UT)	No. of exps.
HD 4306	00 45 27	-09 32 44	9.71	8.76	2400 s	250	19th Oct. 2016	3
HD 6268	01 03 18	-27 52 54	8.89	7.80	1600 s	350	19th Oct. 2016	2

– all the extracted spectral orders were normalised individually before merging.

Then, the spectra of each star were Doppler-shifted to put them in the laboratory reference frame, cleaned by cosmic ray, and, eventually averaged to achieve the highest possible S/N. Using DER_SNR, a ‘simple and general spectroscopic signal-to-noise measurement algorithm’ (Stoehr et al. 2007), we obtained $S/N \approx 390/640$ and $S/N \approx 430/720$, for HD 4306 and HD 6268 respectively, in the blue and in the red parts of the spectra.

3. Atmospheric parameters

We re-determined the stellar parameters of HD 4306 and HD 6268 in order to take advantage of the very high S/N of our mean spectra, even if these stars have been studied several times in the literature; among the most recent results with a detailed chemistry, we used the works by Honda et al. (2004), Roederer et al. (2014), and Mashonkina et al. (2017) for a direct comparison. The full list of sixteen original atmospheric parameter determinations for HD 6268 is presented in Table A.1, with the first determination from 40 yr ago by Luck & Bond (1981), and the most recent one by Arentsen et al. (2019).

We derived the stellar parameters of the two stars via the spectroscopic approach; namely, the effective temperature T_{eff} was determined by requiring the iron abundance from Fe I lines to be independent of their excitation potential, and the surface gravity $\log g$ was derived from the ionisation equilibrium of Fe I and Fe II. The microturbulence velocity ξ was estimated by requiring the $[\text{Fe}/\text{H}]$ derived from the Fe I lines to be independent of their equivalent widths. The equivalent widths of a list of 81 Fe I and of 14 Fe II lines and of 99 Fe I and 16 Fe II lines were measured in the spectra of HD 6268 and HD 4306, respectively. Then, an iterative procedure was adopted to derive the stellar parameters using as initial guess for the parameter values those taken from the latest literature. We used the ABUNDANCE and the BLACKWEL codes⁴ to derive estimates of T_{eff} , $\log g$, $[\text{Fe}/\text{H}]$, and ξ . In Fig. 1, we show that at the adopted temperature $T_{\text{eff}} = 4500$ K the iron abundance of HD 6268 obtained from Fe I lines is almost independent of their excitation potential; on the other hand, a significant slope is obtained at a temperature lower/higher of 100 K. The atmosphere models of the two stars were calculated for the newly derived stellar parameters by using ATLAS12 (Kurucz 2005). As discussed in Kurucz (2005) and Castelli (2005), ATLAS12 can generate an atmosphere model for any desired chemical composition of individual elements and microturbulent velocity, since its treatment of opacity is based on the opacity sampling technique instead of the use of opacity distribution functions as, for example, in ATLAS9.

Table 2 shows the obtained atmospheric parameters for HD 4306 and HD 6268 together with estimates from the literature. We found atmospheric parameters that are reasonably

⁴ <https://www.appstate.edu/~grayro/spectrum/spectrum.html>

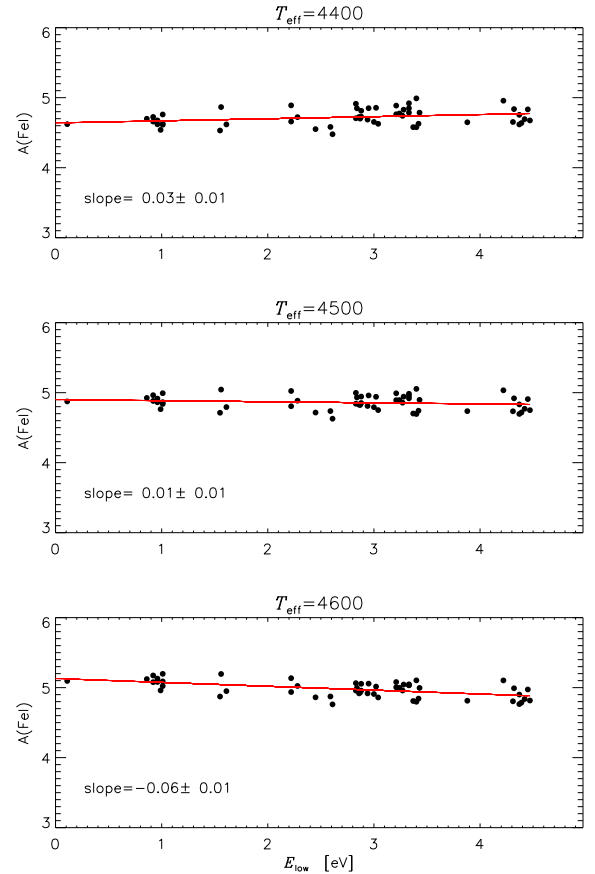


Fig. 1. Iron abundance from Fe I lines versus excitation potential of HD 6268 for the adopted $T_{\text{eff}} = 4500$ K (middle panel) and for a temperature lower ($T_{\text{eff}} = 4400$ K, top panel) and higher ($T_{\text{eff}} = 4600$ K, bottom panel).

close with those available in literature, even if we estimate lower effective temperatures for both stars; in fact, both Honda et al. (2004) and Mashonkina et al. (2017) estimated T_{eff} with a different approach based on the photometric data. On the contrary, in Roederer et al. (2014) T_{eff} is derived with our same approach, by requiring that abundances derived from Fe I lines show no trend with the excitation potential and indeed their estimated T_{eff} for HD 6268 is the closest to our result.

4. Chemical abundances from equivalent widths

We used the measured equivalent widths of other weak lines present in our spectra to derive abundances for the elements listed in Table 3. The errors in our abundances reported in the table are based only on the standard deviation to the mean from the abundances measured in the different lines. These errors are considered as lower limits because we did not compute the impact of the atmospheric parameters on the error budget for all

Table 2. Atmospheric parameters.

HD 4306				
T_{eff} (K)	$\log g$ (dex)	[M/H] (dex)	ξ (km s ⁻¹)	Reference
4700 ± 50	1.80 ± 0.10	-2.90 ± 0.10	1.50 ± 0.13	This paper
4810	1.80	-2.89	1.60	Honda et al. (2004)
4960	2.18	-2.74	1.30	Mashonkina et al. (2017)
HD 6268				
T_{eff} (K)	$\log g$ (dex)	[M/H] (dex)	ξ (km s ⁻¹)	Reference
4500 ± 50	0.80 ± 0.10	-2.50 ± 0.10	2.25 ± 0.13	This paper
4600	1.00	-2.63	2.1	Honda et al. (2004)
4570	0.70	-2.69	1.85	Roederer et al. (2014)

Table 3. Chemical abundances of the two stars and comparison to similar results in literature.

HD 4306							
Elem	This paper			Honda et al. (2004)	Mashonkina et al. (2017)		Roederer et al. (2014)
	[X/H] - LTE	A(X)	N lines	A(X) - LTE	A(X) LTE	A(X) NLTE	A(X) -LTE
Mg I	-2.49 ± 0.09	5.04	4	5.25	5.27	5.30	–
Ca I	-2.49 ± 0.02	3.82	14	3.97	3.98	4.11	–
Sc II	-2.75 ± 0.07	0.42	5	0.45	–	–	–
Ti I	-2.77 ± 0.02	2.13	11	2.55	2.41	2.63	–
Ti II	-2.33 ± 0.04	2.57	23	2.55	2.58	2.59	–
Cr II	-2.74 ± 0.04	2.90	2	2.97	–	–	–
Mn I	-3.67 ± 0.01	1.72	2	2.08	–	–	–
Fe I	-3.00 ± 0.01	4.45	99	4.62	4.60	4.71	–
Fe II	-2.90 ± 0.02	4.55	16	4.62	4.76	4.76	–
Sr II	-3.27	-0.35	1	-0.08	0.17	0.13	–
Ba II	-4.05 ± 0.07	-1.88	2	-1.84	-1.62	-1.62	–
HD 6268							
Elem	This paper			Honda et al. (2004)	Mashonkina et al. (2017)		Roederer et al. (2014)
	[X/H] - LTE	A(X)	N lines	A(X) - LTE	A(X) LTE	A(X) NLTE	[X/H] -LTE
Mg I	-2.13 ± 0.17	5.40	4	5.44	–	–	5.35
Ca I	-2.27 ± 0.02	4.04	13	4.09	–	–	3.98
Sc II	-2.65 ± 0.08	0.52	4	0.51	–	–	0.39
Ti I	-2.55 ± 0.02	2.35	10	2.62	–	–	2.25
Ti II	-2.23 ± 0.04	2.67	14	2.62	–	–	2.54
Cr II	-2.37 ± 0.11	3.27	2	3.20	–	–	3.11
Mn I	-3.10 ± 0.01	2.29	2	2.49	–	–	2.54
Fe I	-2.51 ± 0.01	4.94	81	4.88	–	–	4.68
Fe II	-2.51 ± 0.03	4.94	14	4.88	–	–	4.81
Sr II	-2.59	0.33	1	0.36	–	–	0.33
Ba II	-2.62 ± 0.07	-0.45	2	-0.45	–	–	-0.46

Notes. For solar abundances – [X/H] – we considered [Grevesse et al. \(2007\)](#). For the literature data, we also provide the information on how the abundances were computed assuming the LTE approximation or NLTE.

the elements. Nevertheless, in Table 4 we show the impact of the variation of the atmospheric parameters on the abundance of barium for HD 6268. The variation is small, 0.03 dex, and there is a total budget of 0.08 dex for the errors on the abundance of barium of this star; a similar impact is expected for HD 4306. For Sr, we can measure only one line, and for this reason we cannot compute this error. It is worth noticing that after deriving the abundances of the different elements with the code

ABUNDANCE, we used them to recompute the corresponding atmosphere models with ATLAS12 for the two stars, and we iterated in order to have a full consistency between the abundances derived from the spectra and those used in constructing the atmosphere structures.

Even if a detailed study of the chemical composition of HD 4306 and HD 6268 is not the purpose of this paper, in which we want to study the profiles of Ba II lines, the results shown

Table 4. Variations of barium abundances obtained from the lines at 5853.668 Å and 6141.730 Å as a function of the atmospheric parameters.

	Best model	$\Delta T_{\text{eff}} = 50 \text{ K}$	$\Delta T_{\text{eff}} = -50 \text{ K}$	$\Delta \log(g) = +0.1$	$\Delta \log(g) = -0.1$
[Ba/H] from 5853.668 Å	-2.611	-2.581	-2.639	-2.585	-2.637
[Ba/H] from 6141.730 Å	-2.635	-2.599	-2.669	-2.608	-2.662
mean [Ba/H]	-2.62	-2.59	-2.65	-2.60	-2.65
Δ [Ba/H]		0.03	-0.03	0.02	0.03

Notes. We present the abundance derived from the single barium lines with three digits; in fact, they are considered to compute the mean and Δ [Ba/H] (the following as usual with two digits).

in Table 3 indicate that, starting from our measured equivalent widths and adopting the derived atmosphere parameters, we obtain abundances that are consistent with those reported in literature. We thus provide further confirmation of the reliability of the results reported in Table 2. Regarding the abundances, the enrichment in terms of neutron capture elements in the two stars is quite different: HD 6268 is more abundant in heavy elements compared to HD 4306 and has a ratio of [Sr/Ba] ~ 0 ; with abundances of [Sr/Fe] < 0 and [Ba/Fe] < -1 , HD 4306 presents a high [Sr/Ba] > 0.75 . It is important to confirm this enrichment pattern, since it is on this basis that we select these stars to have, respectively, an expected barium with *r*- or *s*-process isotopic ratios, following the theoretical outcomes of Cescutti & Chiappini (2014). It must be pointed out that the different abundance values we derived for Fe I and Fe II and for Ti I and Ti II are the result of deviations from local thermodynamic equilibrium (LTE), in particular for Fe I and Ti I (see Mashonkina et al. 2017). It is also worth underlining that the impacts of the non-LTE (NLTE) approach to the barium abundances are negligible for HD 4306 according to the results in Mashonkina et al. (2017). On the other hand, the following estimates were also obtained for the star HD 4306 ($T_{\text{eff}} = 4700$; $\log(g) = 1.8$; [Fe/H] = -2.8): for the barium abundance of -1.88 , $\delta\text{NLTE} = +0.14$ for the 6141 line and $\delta\text{NLTE} = +0.21$ for the 4554 line; for the barium abundance of -1.62 , $\delta\text{NLTE} = +0.11$ for the 6141 line and $\delta\text{NLTE} = +0.17$ for the 4554 line; for the barium abundance of -1.32 , $\delta\text{NLTE} = +0.09$ for the 6141 line and $\delta\text{NLTE} = +0.09$ for the 4554 line (Korotin et al. 2015, and Sergey Korotin, priv. comm.). With these estimates of the NLTE impact, we assume an abundance for Ba too high by 0.06–0.07 dex by considering the abundance obtained from the line at 6141 Å for the synthesis of the line at 4554 Å. Indeed, the difference is not negligible, but still within the error. Finally, we also expect NLTE to have a very small impact on Ba for HD 6268, following the results for a similar star in terms of stellar atmosphere and barium abundance again in the Mashonkina et al. (2017) dataset (UMi 36 886).

5. Spectral synthesis

5.1. Synthesis of calcium and iron lines

Using the obtained atmosphere models and the stellar abundances, we computed the LTE synthetic spectra of HD 4306 and of HD 6268 with SPECTRUM. In order to compare them with those observed, we need to take into account the broadening introduced by the instrument and by macroturbulence, V_{mac} , and rotation, $V \sin i$. The shape of the instrumental profile used in degrading the synthetic spectra to the UVES resolution was derived by fitting the profile of several emission lines in the ThAr lamp spectra taken during the observing run. This kind of

analysis provides an instrumental profile, which is a linear combination of a Gaussian and a Lorentzian both with $R = 110\,000$ and with relative weights of 0.9 and 0.1, respectively. As far as macroturbulence and rotation broadening are concerned, we compared the observed and broadened synthetic profiles of three lines: CaI at 6102.7 Å, CaI at 6122.2 Å, and FeII at 4555.8 Å. We used the combination of Gaussian profiles computed with V_{mac} varying from 0 to 6 km s $^{-1}$ at step 0.5 km s $^{-1}$ and rotational profiles⁵ with $V \sin i = 0, 1, 2, 3$, and 4 km s $^{-1}$ as broadening functions. In Figs. 2–4, we show the best resulting synthetic broadened profiles compared to the observed ones, achieved by using $V_{\text{mac}} = 2.5$ and 5.5 km s $^{-1}$ and $V \sin i = 3.0$ and 3.0 km s $^{-1}$ for HD 4306 and HD 6268, respectively.

5.2. Synthesis of barium lines

The same broadening was then adopted to obtain synthetic line profiles for four Ba II lines, namely the lines at 4554.0, 4934.1, 5853.7, and 6141.7 Å. The first two lines' profiles are strongly dependent on the Ba isotope ratios, while the other two are almost independent of them. In computing the theoretical line profiles, we used the *s*- and *r*-process contributions to the five isotopes from Arlandini et al. (1999) to construct pure *s*-process and pure *r*-process Ba II lines. The Ba hfs line list is the one we used for the resonant lines at 4554 Å line and 4934 Å is from McWilliam (1998, Table 1), and it is shown here in Table 5 together with the mass fractions considered for each isotope in the *s*-process and *r*-process cases.

Figures 5–6 show comparisons among the observed profiles (black) and the broadened synthetic ones for the lines at 5853.7 Å and at 6141.7 Å. As mentioned above, these lines do not have strong hfs effects, and therefore they do not show any significant difference in their synthetic profiles between the *s*- and the *r* cases, which appear as a single purple line. The comparison between observed and synthetic lines for 5853.7 Å and at 6141.7 Å is good, although the line at 5853.7 Å for HD 4306 is really weak; these two lines were also used to derive the Ba II abundances given in Table 3. On the other hand, the two lines at 4554.0, 4934.1 Å are sensitive to hfs effects and therefore can be used to infer the role of *s*- and *r*-processes in the two stars. In Figs. 7 and 8, we show the comparison of the observed profiles (black) to the broadened synthetic ones computed for pure *r*- (blue) and *s*-process (red). In the case of HD 4306 the observed profiles are very well matched by the synthetic ones computed assuming a pure *s*-process (see also the small residuals plotted at the bottom of the figures). The observed profiles of the

⁵ See Eq. (18.14) in Gray (2008); that is, flux convolution approximation including the limb-darkening coefficient ϵ (in this work $\epsilon = 0.6$).

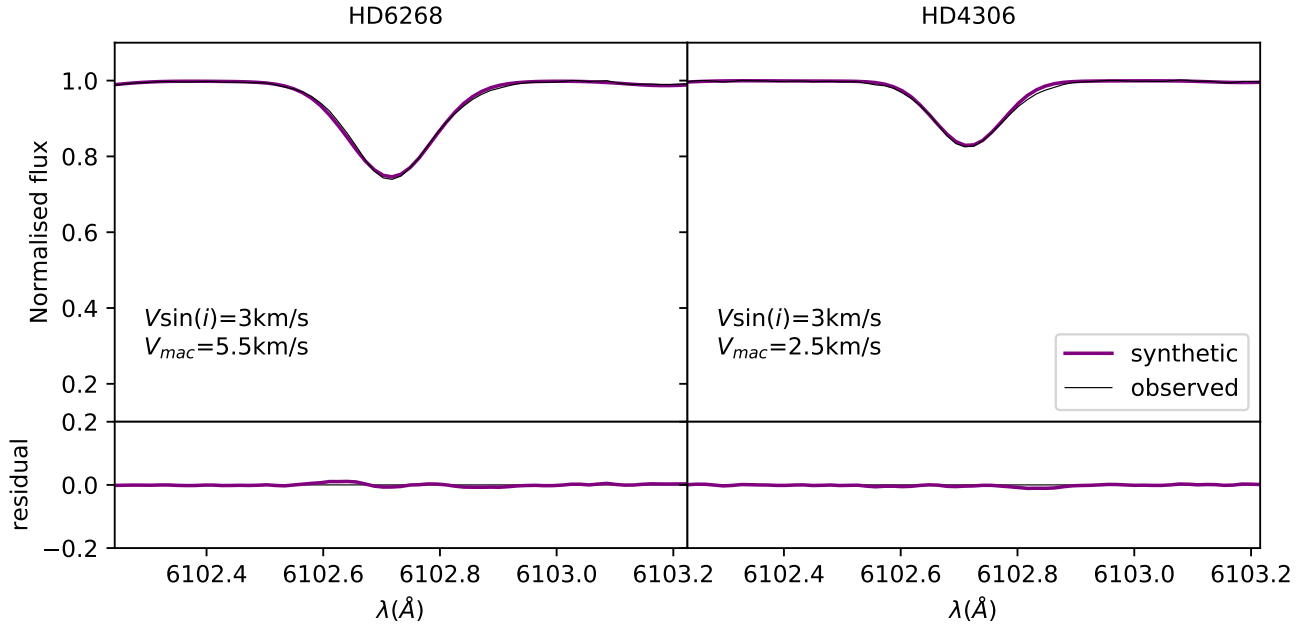


Fig. 2. Ca I line at 6102.7 Å. In the *upper panels*, we show the observed spectra with a black line. The purple lines are the synthetic spectra obtained; in the *lower panels*, the residuals between synthetic spectra and observed spectra are reported. In the *left panels*, we present the results for HD 6268, and on the *right* we show those for HD 4306.

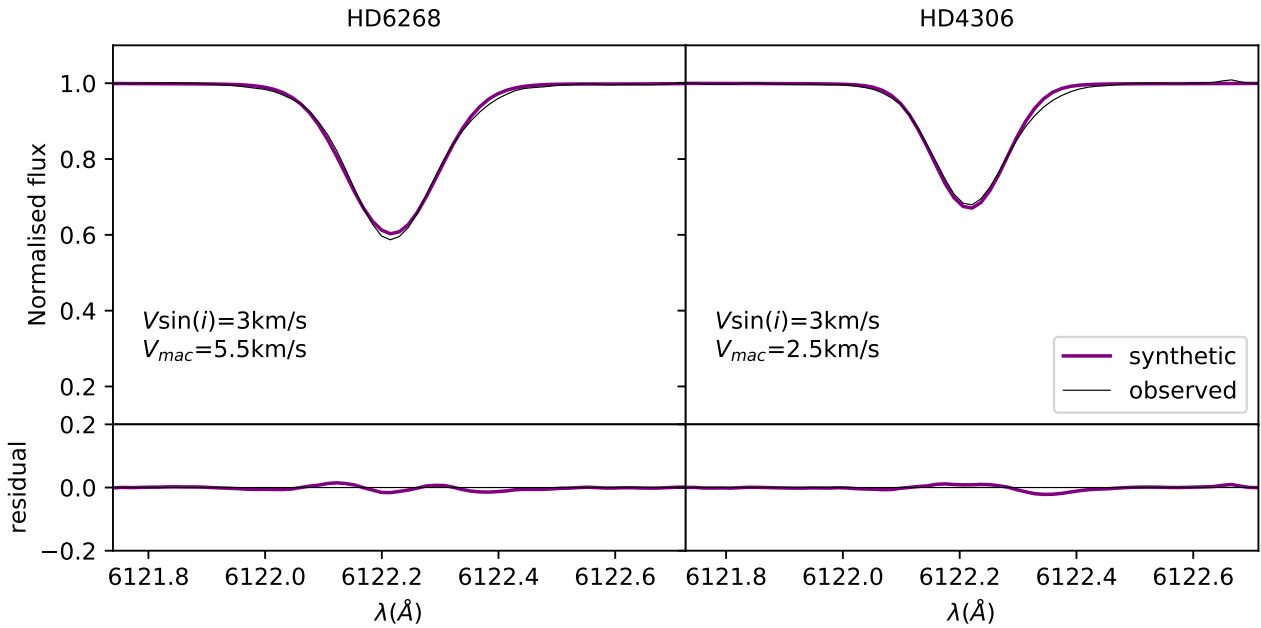


Fig. 3. Same as Fig. 2, but for Ca I line at 6122.2 Å.

4554.0 and 4934.1 Å lines in HD 6268 are not very well reproduced by the synthetic ones, either assuming a pure *s*- or pure *r*-process. On the other hand, the discrepancies are smaller in the case of the pure *r*-process (see also the residuals). In HD 6268, both these two lines are much stronger than the corresponding ones in HD 4306, thus the mismatch of the total equivalent width between observation and synthetic profile can be ascribed to a problem with the adopted ξ or with deviation from LTE and can be explained only in part by differences in the isotopic composition. An increase of the Ba abundance of 0.7 and 0.5 dex is needed to match the observed equivalent width, assuming an *s*- and *r*-process isotopic composition, respectively. The necessary correction is smaller for the *r*-process case, and the resulting line

profile fit is better than that of the *s*-process assumption. This gives additional support for the conclusion that the barium of HD 6268 most likely has an *r*-process origin. In conclusion, by looking at Figs. 7 and 8 we can say that while the Ba II lines in HD 4306 are fully consistent with a pure *s*-process origin, those in HD 6268 are better reproduced assuming the predominance of the *r*-process.

6. Comparison of observed spectra

The results presented thus far are indeed model dependent. In particular, they depend heavily on the accuracy of the hfs data used to build our synthetic spectra and on the reliability of the

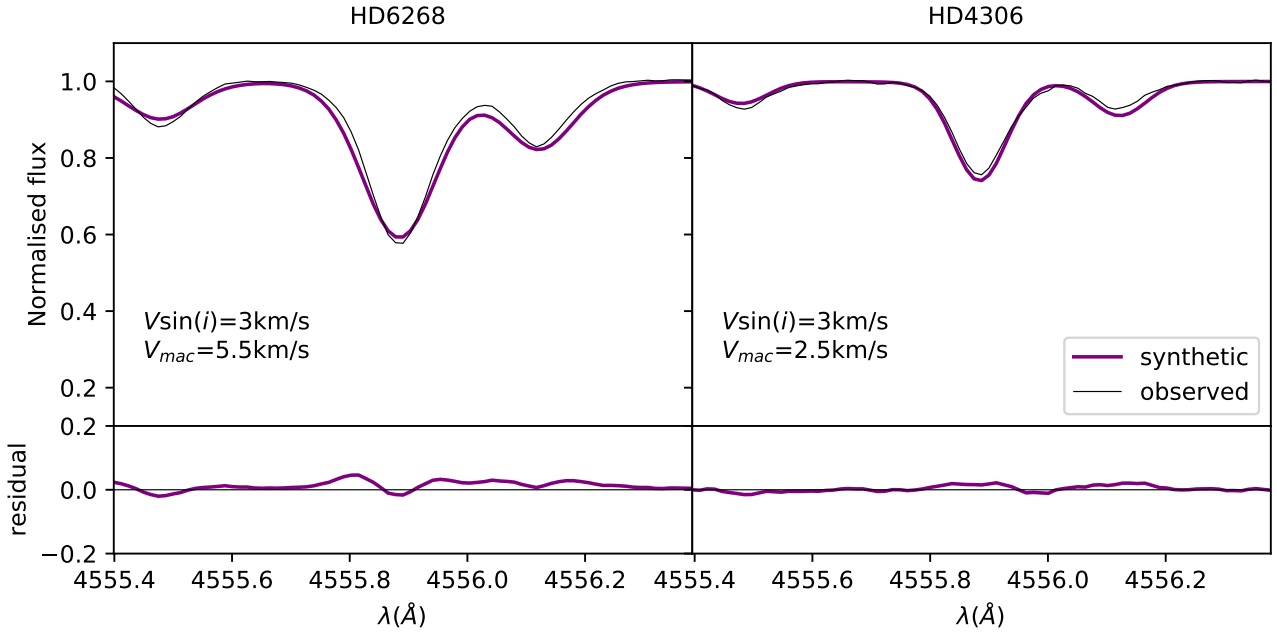


Fig. 4. Same as Fig. 2, but for Fe II line at 4555.8 Å.

Table 5. Parameters of barium lines at 4554 Å and 4934 Å used in our calculations with hfs, including the wavelengths, excitation potentials of the lower level, numbers of isotopes, the mass fraction for each isotope, and oscillator strengths.

Ba II	$E_{\text{low}} = 0.00 \text{ eV}$	$\log gf = +0.17$	$E_{\text{low}} = 0.00 \text{ eV}$	$\log gf = -0.15$	s -process	r -process
	Wavelength (Å)	Strength	Wavelength (Å)	Strength	fraction	fraction
^{134}Ba	4554.032	1.000	4934.075	1.000	0.02	0.00
^{135}Ba	4554.001	0.1562	4934.034	0.3125	0.03	0.40
	4554.002	0.1562	4934.045	0.0625		
	4554.003	0.0625	4934.093	0.3125		
	4554.049	0.4375	4934.104	0.3125		
	4554.052	0.1562				
	4554.053	0.0313				
^{136}Ba	4554.032	1.000	4934.075	1.000	0.10	0.00
^{137}Ba	4553.997	0.1562	4934.029	0.3125	0.09	0.32
	4553.999	0.1562	4934.041	0.0625		
	4554.000	0.0625	4934.096	0.3125		
	4554.052	0.4375	4934.107	0.3125		
	4554.054	0.1562				
	4554.055	0.0313				
^{138}Ba	4554.032	1.0000	4934.075	1.0	0.76	0.28

spectrum and atmosphere model codes in reproducing the real spectra and atmosphere of the two stars. For this reason, it would be worthwhile to perform a relative analysis exclusively using the two observed spectra. Unfortunately, such an approach would require the two stars to be perfectly equal, apart from the relative contributions of the s - and r -process to the formation of Ba. Even if the 4554.0 and 4934.1 Å Ba line profiles depend on the Ba isotope ratios, their profiles are mainly determined by the values of the stellar atmospheric parameters, by the Ba abundance,

and by the ξ , V_{mac} , and $V \sin i$ velocities. In the following, we use the pure- s synthetic spectra of the two stars in a differential way in order to transform the observed spectrum of HD 4306 into the spectrum this star would have if its atmospheric parameters and velocities were exactly the same as those of HD 6268. In such a way, our results are still built with the help of the synthetic spectra, but they should be less model dependent than those described in Sect. 5. The idea is to build a spectrum of a pure- s Ba star with the same stellar atmosphere parameters and

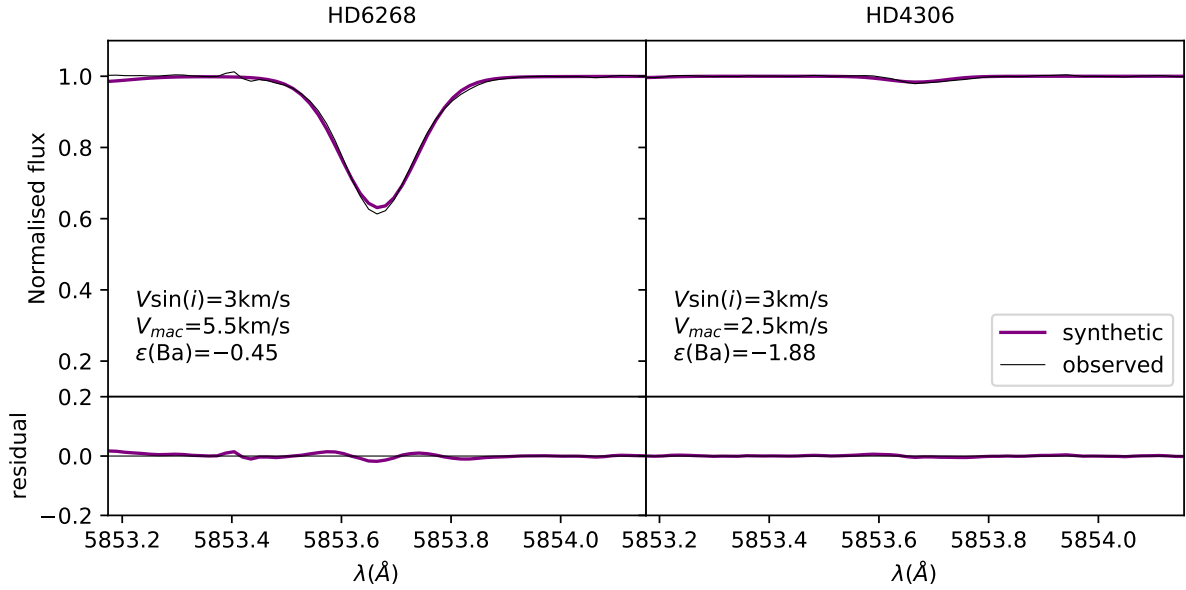


Fig. 5. Ba II line at 5853 Å. In the *upper panels*, we show the observed spectra with a black line. The purple line is actually two nearly identical profiles, computed for pure *s*- and pure *r*-process hfs, respectively; in the *lower panels*, the residuals between synthetic spectra and observed spectra are reported. In the *left panels*, we present the results for HD 6268, and on the *right* we show those for HD 4306.

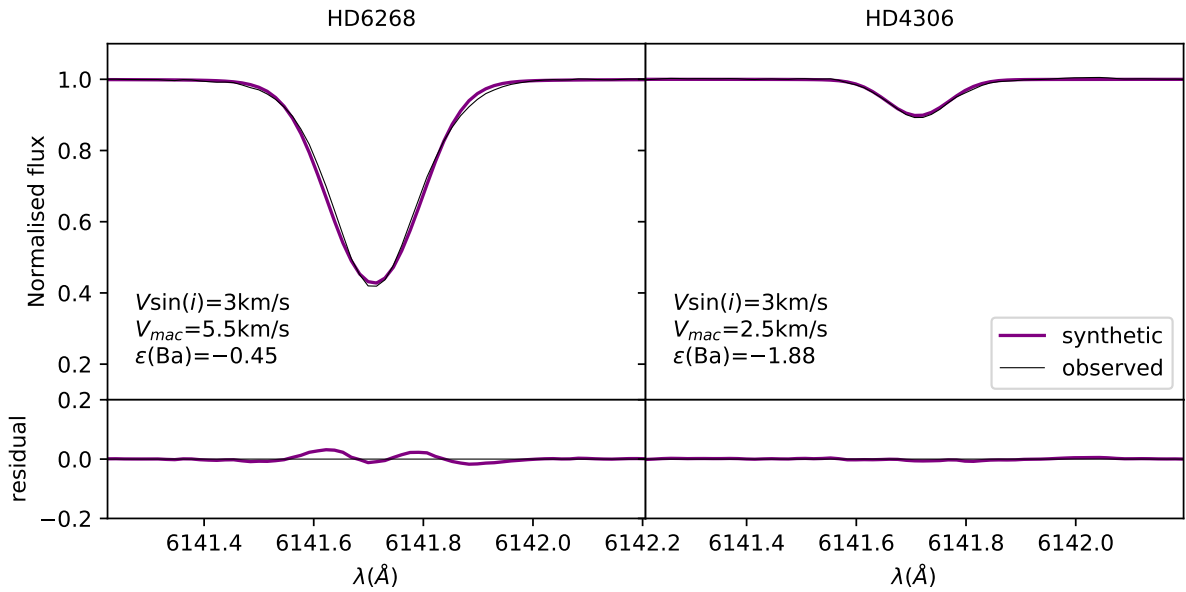


Fig. 6. Same as Fig. 5, but for Ba II line at 6141 Å.

abundances of HD 6268 (hereafter fake-HD 6268) starting from the spectrum of HD 4306 and to compare it with the spectrum of the real HD 6268 in order to search for a possible contribution of non-*s*-processed Ba in the latter star. The lines in HD 4306 and in the fake-HD 6268 spectrum will differ in strength (equivalent width) and broadening due to the different stellar atmosphere characteristics and chemical compositions. If the line profiles are mainly Doppler profiles, it is easy to scale those of HD 4306 (in a velocity scale centred at the nominal wavelength of each line) to obtain the broadening of the fake-HD6268 by taking into account the different effective temperatures and macroturbulence velocities in the two stars. On the other hand, if other broadening mechanisms are present (such as, for example, collision broadening, van der Waals broadening, and Stark broadening) it is not easy to take them into account within a simple equation; therefore we adopted a parametric multiplication factor, b , to address

their effects. Thus, we initially changed from the wavelength scale to a velocity scale; that is, $\Delta v = c \times \frac{(\lambda - \lambda_0)}{\lambda_0}$. Then, to take into account the differences in broadening, we scaled each pure-*s* synthetic HD 4306 line profile to build the corresponding line of the fake-HD 6268 using the following equation:

$$\Delta v^{\text{fake-HD 6268}} = b \cdot \Delta v^{\text{HD 4306}} \sqrt{\frac{T_{\text{eff}}^{\text{HD6268}} V_{\text{mac}}^{\text{HD6268}}}{T_{\text{eff}}^{\text{HD4306}} V_{\text{mac}}^{\text{HD4306}}}} \quad (1)$$

All the resulting fake-HD 6268 lines were then scaled to have the same equivalent width of the corresponding HD 6268 lines removing the effects of the different abundances and macroturbulences. The b value for each line was empirically derived by applying the above procedure to the synthetic pure-*s* spectra of the two stars by iterating and varying b until we achieved a good

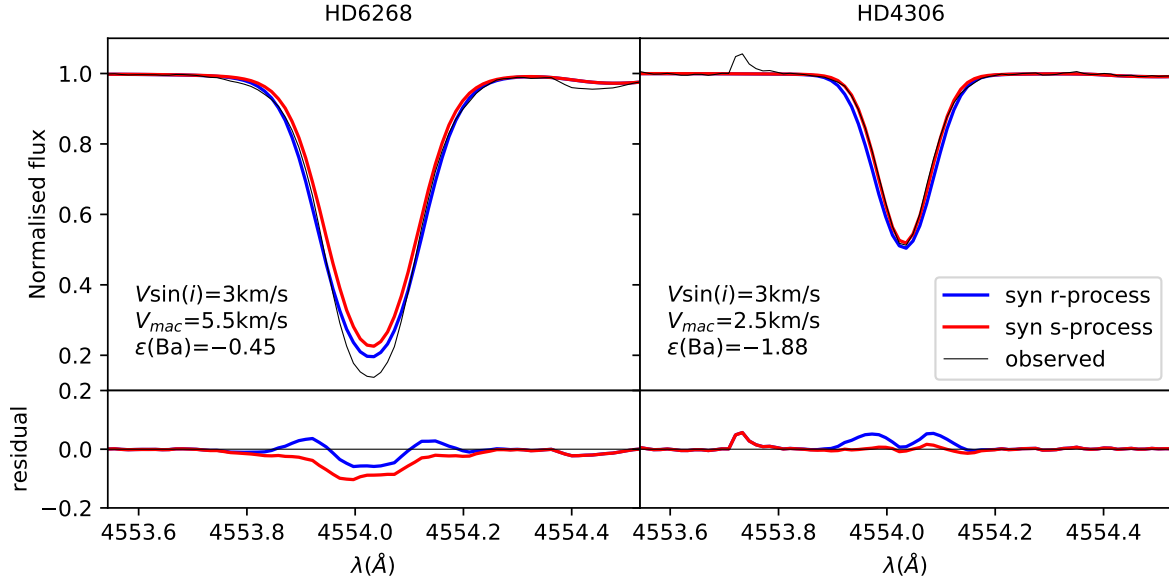


Fig. 7. Ba II line at 4554 Å. In the *upper panels*, we show the observed spectra with a black line. The blue lines are the synthetic spectra obtained with an *r*-process isotopic composition of barium. The red lines are the synthetic spectra obtained with an *s*-process composition. In the *lower panels*, the residual for both synthetic spectra are reported. In the *left panels*, we present the results for HD 6268, and on the *right* we show those for HD 4306.

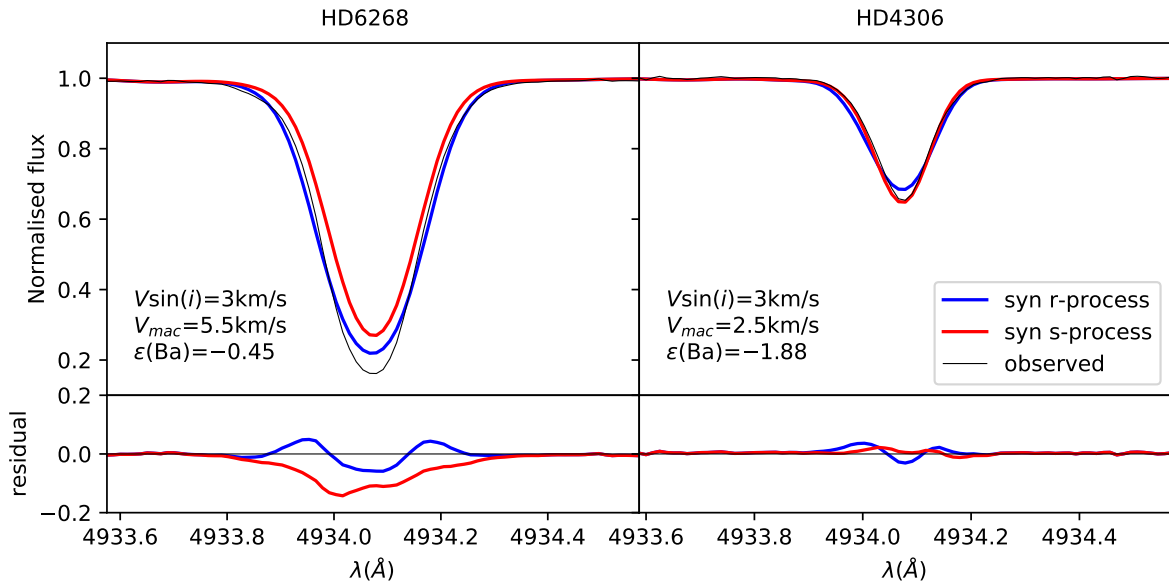


Fig. 8. Same as Fig. 7, but for Ba II line at 4934 Å.

match between the synthetic fake-HD 6268 and the pure-*s* synthetic HD 6268 lines. We did not use the observed spectra, since we wanted to avoid incorporating any broadening in the empirically derived *b* values caused by hfs effects produced by odd isotopes in the Ba lines. In the left panels of Fig. 9, we plot the obtained synthetic line profiles of the fake-HD 6268 (dark green lines) superimposed on the original pure-*s* synthetic profiles of HD 6268 (red lines) after re-sampling all of them at steps of 0.1 km s⁻¹ for all the lines analysed in Sect. 5. This was done with the exception of the Fe II line that is blended with the Cr I line (4556.163 Å) and the Ba II 5853.7 Å line that is too weak in HD 4306. The values of the *b* factor are around one (i.e. 0.97, 0.97, 1.15, 1.35, and 1.27 for Ca I 6102.7, Ca I 6122.2, Ba II 6141.7, Ba II 4554.0, and Ba II 4934.1 Å, respectively), as expected since we already corrected for the Doppler effects. As can be seen in the left column panels of Fig. 9, the fake synthetic profiles

match very well those of HD 6268 apart for the blue wing of the Ba II 4934.1 Å line that has a known blend with a weak Fe I line (see Fig. 5 in Gallagher et al. 2010) since the two stars have different iron abundances. Thus, our procedure works effectively to remove practically all the differences between the two stars and to transform the synthetic pure-*s* profiles of HD 4306 in those we should expect if HD 4306 were actually identical to HD 6268.

Eventually, we apply the so derived individual *b* factors to scale the observed HD 4306 lines to obtain the observed fake-HD 6268 lines by assuming a pure *s*-process for HD 6268. In the right column panels of Fig. 9, we plot the observed fake-HD 6268 profiles (light green lines) superimposed on the corresponding observed profiles of HD 6268 (black lines) in order to check our assumption. Our procedure removed practically all the differences between the two stars and transform

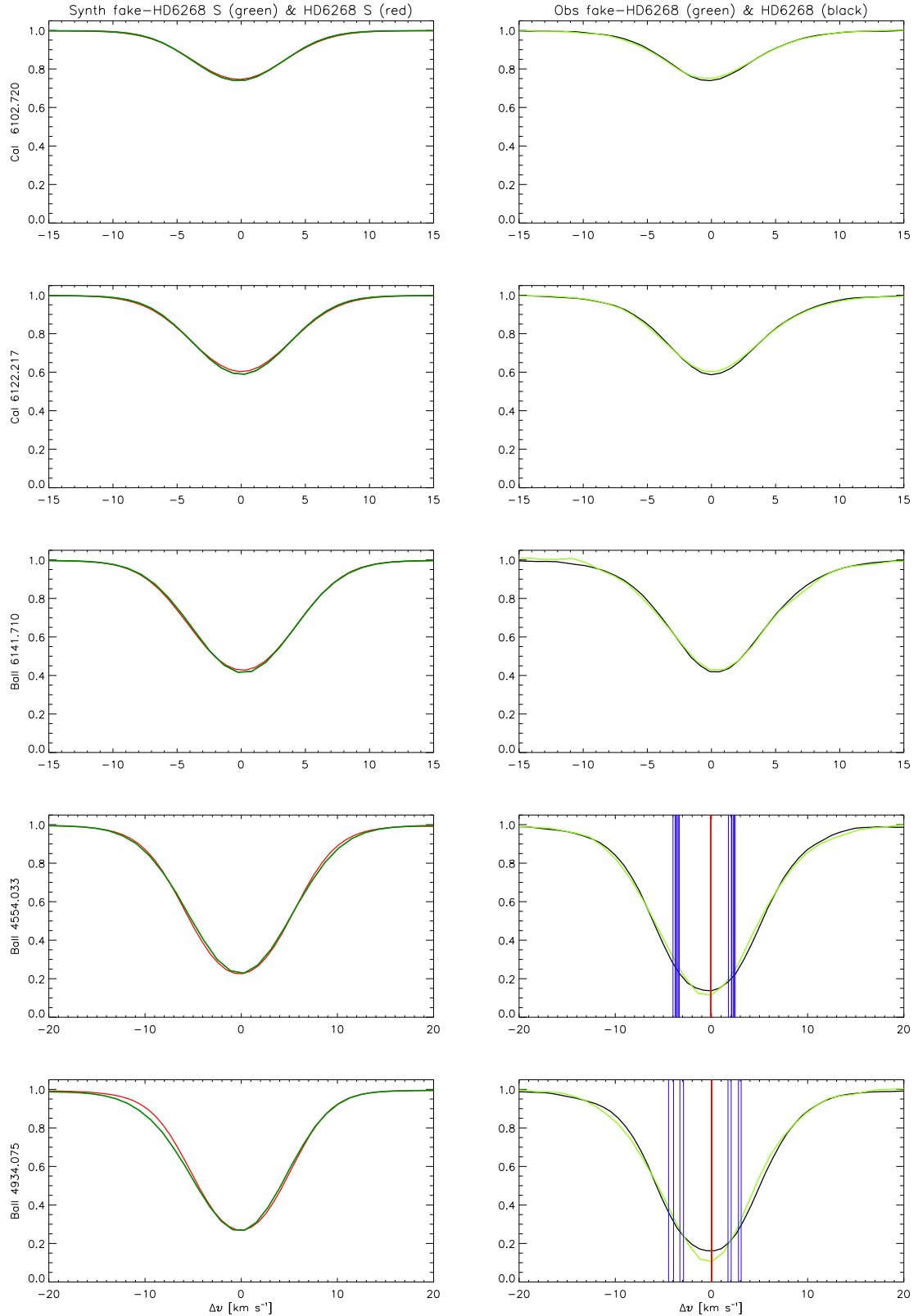


Fig. 9. *Left panels:* comparison of synthetic pure- s profiles of a scaled fake-HD 6268 star (dark green) with those of HD 6268 (red). *Right panels:* comparison of scaled ‘observed’ profiles of a fake-HD 6268 (light green) with the observed ones of HD 6268 (black) (see text). Vertical blue lines show the hfs lines of odd Ba isotopes, and vertical red lines show lines of the even Ba isotopes.

the observed profiles of HD 4306 in those we should expect if HD 4306 was actually identical to HD 6268 when we look at the two Ca I lines and at the Ba II 6141.7 line, which are unaffected by the hypothesis that both stars are pure- s Ba objects. On the

other hand, when we apply the same technique to the 4554.0 and 4934.1 Å Ba II lines, whose profiles are sensitive to the relative contributions of s - and r -processes in the two stars, the observed fake-HD 6268 profiles do not reproduce the observed profiles of

HD 6268 as well as in the case of the corresponding synthetic ones (left panels in the bottom two lines of Fig. 9). In particular, we notice main differences in the cores where the even Ba isotopes shape the lines and in the regions of the hfs lines of odd Ba isotopes, indicated by vertical red and blue lines in Fig. 9, respectively. So, these results suggest that the assumption of a pure s -process for HD 6268 should be rejected since it leads to fake observed profiles of HD 6268 that are different from the real observed ones, in particular in the region sensitive to s - and r -process contribution.

We cannot firmly conclude that the barium isotopes in HD 6268 have a pure r -process pattern. On the other hand, the differences shown in the right panels of the two bottom lines should be due to an effect that does not occur in the other lines. In this respect, the most likely reason is the different pattern of barium isotopes between the two stars; therefore, these differences can be ascribed to the enrichment of an r -process with a large fraction of odd isotopes in HD 6268, absent from the fake-HD 6268 spectrum. This outcome arises from the differential analysis of synthetic and observed spectra, and so it is almost independent of the theoretical hfs Ba data; the only dependence is in the negligible hfs effects of the scarce odd isotopes within the s -process pattern.

7. Conclusions

According to our results, the two stars analysed here present two different isotopic compositions of barium. In fact, the results obtained in Sects. 5 and 6 are both consistent with the hypothesis that the barium in HD 4306 was produced exclusively by an s -process, while the one in HD 6268 also shows a contribution by an r -process. The predictions of the previous chemical evolution models in Cescutti & Chiappini (2014) pointed out exactly this possibility, that in the Galactic halo both an s -process and an r -process can pollute the ISM and therefore preserve this differential signature. The high [Ba/Fe] stars that typically show a solar or slightly sub-solar [Sr/Ba] ratio are most likely polluted by an r -process source, in this case HD 6268 with [Ba/Fe] = 0.09 and [Sr/Ba] = 0.03. On the other hand, an s -process source should have polluted the stars with a low [Ba/Fe] and with a [Sr/Ba] that, according to the theoretical yields for rotating massive stars, can go from about solar (possibly even below solar) up to a [Sr/Ba] \sim 1.5 (Frischknecht et al. 2016; Limongi & Chieffi 2018), in our case HD 4306 with [Ba/Fe] = -1.05 and [Sr/Ba] = 0.78. Therefore, it could be that the light elements' primary production, empirically introduced by Travaglio et al. (2004), is an s -process that is activated almost as a primary process by rotation. This possibility was not excluded by the study of Montes et al. (2007), together with the more commonly invoked weak r -process. Clearly, the fact that this happens for two stars does not provide firm conclusions, and more stars must be measured to determine the real statistical impact of the s -process in massive stars at low metallicity. Nevertheless, we think that this work provides insight into the impact of detailed studies of isotopes and the possibilities that high-quality spectra can provide.

Acknowledgements. MV and NC acknowledge support by Deutsche Forschungsgemeinschaft (DFG, German Research Foundation) – project-IDs: 428473034, and 138713538 (SFB 881 “The Milky Way System”, subproject A04). CK

acknowledges funding from the UK Science and Technology Facility Council (STFC) through grant ST/M000958/1 & ST/V000632/1. This work was partially supported by the European Union (ChETEC-INFRA, project no. 101008324 and ChETEC, CA16117).

References

- Arcones, A., & Montes, F. 2011, *ApJ*, 731, 5
- Arentsen, A., Prugniel, P., Gonneau, A., et al. 2019, *A&A*, 627, A138
- Arlandini, C., Käppeler, F., Wisshak, K., et al. 1999, *ApJ*, 525, 886
- Barbuy, B., Spite, F., & Spite, M. 1985, *A&A*, 144, 343
- Barklem, P. S., Christlieb, N., Beers, T. C., et al. 2005, *A&A*, 439, 129
- Castelli, F. 2005, *Mem. Soc. Astron. Ital. Suppl.*, 8, 34
- Cenarro, A. J., Peletier, R. F., Sánchez-Blázquez, P., et al. 2007, *MNRAS*, 374, 664
- Cescutti, G., & Chiappini, C. 2014, *A&A*, 565, A51
- Cescutti, G., Chiappini, C., Hirschi, R., Meynet, G., & Frischknecht, U. 2013, *A&A*, 553, A51
- Chiappini, C., Hirschi, R., Meynet, G., et al. 2006, *A&A*, 449, L27
- Chiappini, C., Ekström, S., Meynet, G., et al. 2008, *A&A*, 479, L9
- Chiappini, C., Frischknecht, U., Meynet, G., et al. 2011, *Nature*, 472, 454
- Chopin, A., Hirschi, R., Meynet, G., et al. 2018, *A&A*, 618, A133
- Collet, R., Asplund, M., & Nissen, P. E. 2009, *PASA*, 26, 330
- Dekker, H., D’Odorico, S., Kaufer, A., Delabre, B., & Kotzlowski, H. 2000, in *SPIE Conf. Ser.*, 4008, *Optical and IR Telescope Instrumentation and Detectors*, eds. M. Iye, & A. F. Moorwood, 534
- Francois, P. 1996, *A&A*, 313, 229
- François, P., Depagne, E., Hill, V., et al. 2007, *A&A*, 476, 935
- Frischknecht, U., Hirschi, R., Pignatari, M., et al. 2016, *MNRAS*, 456, 1803
- Gallagher, A. J., Ryan, S. G., García Pérez, A. E., & Aoki, W. 2010, *A&A*, 523, A24
- Gallagher, A. J., Ryan, S. G., Hosford, A., et al. 2012, *A&A*, 538, A118
- Gallagher, A. J., Ludwig, H. G., Ryan, S. G., & Aoki, W. 2015, *A&A*, 579, A94
- Gratton, R. G. 1989, *A&A*, 208, 171
- Gray, D. F. 2008, *The Observation and Analysis of Stellar Photospheres*
- Grevesse, N., Asplund, M., & Sauval, A. J. 2007, *Space Sci. Rev.*, 130, 105
- Honda, S., Aoki, W., Kajino, T., et al. 2004, *ApJ*, 607, 474
- Jablonka, P., North, P., Mashonkina, L., et al. 2015, *A&A*, 583, A67
- Kobayashi, C., Karakas, A. I., & Lugaro, M. 2020, *ApJ*, 900, 179
- Korotin, S. A., Andrievsky, S. M., Hansen, C. J., et al. 2015, *A&A*, 581, A70
- Kurucz, R. L. 2005, *Mem. Soc. Astron. Ital. Suppl.*, 8, 14
- Lambert, D. L., & Allende Prieto, C. 2002, *MNRAS*, 335, 325
- Limongi, M., & Chieffi, A. 2018, *ApJS*, 237, 13
- Luck, R. E., & Bond, H. E. 1981, *ApJ*, 244, 919
- Luck, R. E., & Bond, H. E. 1985, *ApJ*, 292, 559
- Magain, P. 1995, *A&A*, 297, 686
- Mashonkina, L., Jablonka, P., Sitnova, T., Pakhomov, Y., & North, P. 2017, *A&A*, 608, A89
- McDonald, L., Zijlstra, A. A., & Watson, R. A. 2017, *MNRAS*, 471, 770
- McWilliam, A. 1998, *AJ*, 115, 1640
- McWilliam, A., Preston, G. W., Sneden, C., & Searle, L. 1995, *AJ*, 109, 2757
- Meléndez, J., & Barbuy, B. 2002, *ApJ*, 575, 474
- Montes, F., Beers, T. C., Cowan, J., et al. 2007, *ApJ*, 671, 1685
- Pilachowski, C. A., Sneden, C., & Kraft, R. P. 1996, *AJ*, 111, 1689
- Prantzos, N. 2012, *A&A*, 542, A67
- Prantzos, N., Abia, C., Limongi, M., Chieffi, A., & Cristallo, S. 2018, *MNRAS*, 476, 3432
- Prugniel, P., Vauglin, I., & Koleva, M. 2011, *A&A*, 531, A165
- Raiteri, C. M., Gallino, R., & Busso, M. 1992, *ApJ*, 387, 263
- Rizzuti, F., Cescutti, G., Matteucci, F., et al. 2019, *MNRAS*, 489, 5244
- Rizzuti, F., Cescutti, G., Matteucci, F., et al. 2021, *MNRAS*, 502, 2495
- Roederer, I. U., Preston, G. W., Thompson, I. B., et al. 2014, *AJ*, 147, 136
- Rutten, R. J. 1978, *Sol. Phys.*, 56, 237
- Sharma, K., Prugniel, P., & Singh, H. P. 2016, *A&A*, 585, A64
- Spite, F., Spite, M., Barbuy, B., et al. 2018, *A&A*, 611, A30
- Stoehr, F., Fraquelli, D., Kamp, I., et al. 2007, *Space Telescope Eur. Coord. Facility Newsl.*, 42, 4
- Thielemann, F.-K., Arcones, A., Käppeli, R., et al. 2011, *Progr. Particle Nuclear Phys.*, 66, 346
- Travaglio, C., Gallino, R., Arnone, E., et al. 2004, *ApJ*, 601, 864
- Wu, X., Wang, L., Shi, J., Zhao, G., & Grupp, F. 2015, *A&A*, 579, A8

Appendix A: Additional Table

Table A.1. List of the atmosphere parameters of HD 6268 from literature. Column ‘r’ indicates the data type: ‘H’ is high-resolution spectra, ‘L’ is low-resolution spectra, and ‘P’ is photometry. Column ‘c’ indicates whether the data are original measurements (‘O’) or a compilation of re-calibrated sources (‘C’). Column ‘m’ indicates the method: ‘A’ stands for model atmosphere and ‘T’ stands for empirical calibration.

Teff	err	logg	err	[Fe/H]	err	r	c	m	Reference
4667		1.25		-2.42		H	O	A	Luck & Bond (1981)
4755		1.2		-2.1		H	O	A	Luck & Bond (1985)
4582		0.75		-2.4		H	O	A	Barbuy et al. (1985)
4800	100	0.84	0.03	-2.56	0.05	H	O	A	Gratton (1989)
4670		0.75		-2.58		H	O	A	McWilliam et al. (1995)
4800	100	0.84	0.03	-2.52	0.05	H	O	A	Francois (1996)
4700		1.6		-2.36		H	O	A	Pilachowski et al. (1996)
4705		1.5		-2.35		H	O	A	Meléndez & Barbuy (2002)
4600		1.		-2.62	0.11	H	O	A	Honda et al. (2004)
4740		1.2		-2.32			C		Cenarro et al. (2007)
4735	113	1.42	0.26	-2.36	0.11	L	O	T	Prugniel et al. (2011)
4570	34	0.7	0.16	-2.82	0.15	H	O	A	Roederer et al. (2014)
4726		1.14		-2.63		H	O	A	Wu et al. (2015)
4571	161	1.13	0.39	-2.63	0.18	L	O	T	Sharma et al. (2016)
4837	125	1.59				P	O	T	McDonald et al. (2017)
4662	47	1.37	0.14	-2.47	0.07	H	O	A	Arentsen et al. (2019)