












# Multiwavelength Photometry and Progenitor Analysis of the Nova V906 Car

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## Abstract

We present optical and infrared photometry of the classical nova V906 Car, also known as Nova Car 2018 and ASASSN-18fv, which was discovered by the All-Sky Automated Survey for SuperNovae (ASAS-SN) on 2018 March 16.32 UT (MJD 58193.0). The nova reached its maximum on MJD 58222.56 at  $V_{\max} = 5.84 \pm 0.09$  mag, and had decline times of  $t_{2,V} = 26.2$  days and  $t_{3,V} = 33.0$  days. The data from Evryscope shows that the nova had already brightened to  $g' \simeq 13$  mag five days before discovery, as compared with its quiescent magnitude of  $g = 20.13 \pm 0.03$ . The extinction toward the nova, as derived from high-resolution spectroscopy, shows an estimate consistent with foreground extinction to the Carina Nebula of  $A_V = 1.11^{+0.54}_{-0.39}$ . The light curve resembles a rare C (cusp) class nova with a steep decline slope of  $\alpha = -3.94$  post-cusp flare. From the light-curve decline rate, we estimate the mass of the white dwarf to be  $M_{\text{WD}} = < 0.8 M_{\odot}$ , consistent with  $M_{\text{WD}} = 0.71^{+0.23}_{-0.19}$  derived from modeling the accretion disk of the system in quiescence. The donor star is likely a K-M dwarf of  $0.23\text{--}0.43 M_{\odot}$ , which is heated by its companion.

*Unified Astronomy Thesaurus concepts:* Classical novae (251); Extinction (505); Photometry (1234)

*Supporting material:* data behind figures

## 1. Introduction

Eruptions of classical novae (CNe) occur on the surface of a mass-accreting white dwarf (WD) from a hydrogen-rich brown dwarf, red dwarf, red giant companion, or a helium star in a close binary system (Warner 1995). When sufficient mass is accumulated to the point where the degenerate electron pressure at the base of the mass envelope exceeds a critical value, a thermonuclear runaway occurs (Starrfield et al. 2016). Most or all of the envelope is ejected, and the luminosity increases up to or even beyond the WD Eddington luminosity (Shara 1981). The luminosity of a nova eruption is typically in excess of  $10^5 L_{\odot}$ , making CNe among the most luminous stars in the Galaxy (Shara et al. 2016).

Historically, due to the lack of coverage, few novae have been caught on their initial rise (Bode & Evans 2008), and it was only possible to construct parts of their early rise from patrol images of the sky taken by amateur astronomers (see, for example, V1500 Cyg Liller et al. 1975). However, with increasingly better transient event detectors, such as the All-Sky Survey for SuperNovae (ASAS-SN; Shappee et al. 2014) and Evryscope (Law et al. 2016) among others, more CNe have been caught on these cameras on their early rise, allowing astronomers to gather precious early-time data of novae outbursts. These observations are crucial, as they provide insights on many previously poorly observed and understood phases of novae evolution (Hounsell et al. 2016).

CNe can be classified through various properties, such as their orbital periods (Patterson 1984), magnetic strength (Warner 1995), and spectra (Williams et al. 1991). Another popular method of

classifying CNe is through their light-curve shape. Based on a statistical analysis by Stroepe et al. (2010), a majority of CNe are S class CNe, having a smooth, power-law decline with no major fluctuation. However, there are many peculiar novae that do not conform to this type of evolution. One such type is the C class nova, which has a characteristic secondary maximum after the primary peak and appears to be a substantial superimposed brightening above the base level power-law decline (Stroepe et al. 2010).

With the increased prevalence of spectral data of CNe, astronomers have found that there is a statistical correlation between decline rate and the mass of the progenitor WD. Hachisu & Kato (2006) developed an optical and near-infrared light-curve model of novae in which bremsstrahlung from optically thin ejecta dominates the continuum flux. They derived a “universal decline law” and introduced a “timescaling factor” to their light-curve template. Using this timescaling factor, Hachisu & Kato (2006) were able to estimate various nova properties, such as the optically thick wind phase and the WD mass.

In this paper, we present the results of our photometric follow-up campaign of the nova V906 Car (also known as ASASSN-18fv and Nova Car 2018) and its progenitor system. The nova was discovered (Stanek et al. 2018) by ASAS-SN (Shappee et al. 2014) on 2018 March 20.32 UT ( $T_D = 58197$  MJD), which we adopt as our reference epoch. The transient was located at R.A. =  $10^{\text{h}}36^{\text{m}}15^{\text{s}}.43$  and decl. =  $-59^{\circ}35'53''.73$  J2000. This position is updated using the Gaia DR2 data (Gaia Collaboration et al. 2018) available for the progenitor,

corrected to J2000. The nova is located near the Galactic plane, in the same field of view as the Carina Nebula. After discovery, multiple optical and near-infrared spectroscopic observations confirmed V906 Car to be a CN a few days after its discovery (Izzo et al. 2018; Luckas 2018; Rabus & Prieto 2018). The nova was also serendipitously observed by the BRiGht Target Explorer (BRITE) Constellation (Weiss et al. 2014), while monitoring the nearby red giant star HD 92063. The light curve, which contains the full rise of the nova, shows that the outburst started at about six days before the date of discovery.<sup>8</sup>

In a recent study of Car V906, Aydi et al. (2020) analyzed high cadence concurrent  $\gamma$ -ray and optical data taken during the nova outburst. The high correlation between the flaring emission in both wavelengths suggested that the bulk of the nova luminosity was shock powered. This challenges the traditional picture, where the nova luminosity originates in sustained nuclear burning on the surface of a WD after the initial eruption. The study also found that the X-ray emission in Car V906 was highly attenuated, consistent with highly embedded shocks in the ejecta (Nelson et al. 2018). The shock-powered X-ray emission was reprocessed, and emerged at longer wavelengths.

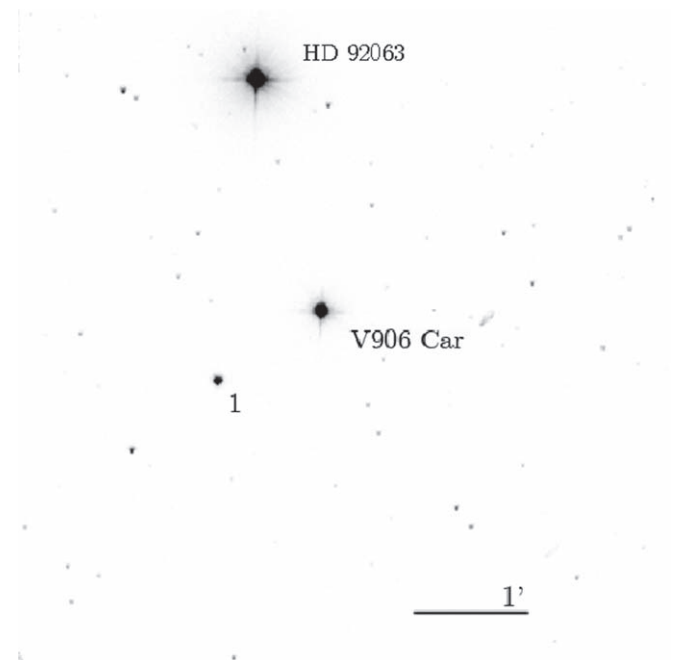
Here we focus on the observational properties of V906 Car and provide the first analysis of the system in quiescence. We present our follow-up observations of the nova in Section 2, spanning from six days before discovery, up to a year after. We estimate the extinction toward the nova in Section 3 and describe the nova photometric evolution in Section 4. We analyze the nova light curve and the progenitor system in Section 5. In Section 6 we discuss the classification of the nova and estimate its WD mass based on decline rates. We summarize our conclusions in Section 7.

## 2. Observations

### 2.1. CTIO 1.3 m ANDICAM

We used the Yale SMARTS 1.3 m Telescope at the Cerro Tololo Inter-American Observatory (CTIO) to obtain optical and near-infrared (near-IR) observations in *BVR/IR* bandpasses. The images were taken under the program YNUS-16A-0001 (PI: B. Penprase) with the ANDICAM instrument, an imager permanently mounted on the 1.3 m telescope that takes simultaneous optical and infrared data. The ANDICAM instrument uses the standard Johnson *BV* filters, Kron-Cousins *RI* filters, and standard CIT/CTIO *JHK* filters. The optical field of view is  $6'.3 \times 6'.3$ , while the infrared (IR) field of view is  $2'.34 \times 2'.34$ . With the CCD readout in  $2 \times 2$  binning mode, the ANDICAM instrument gives a plate scale on the 1.3 m telescope of  $0''.369 \text{ pixel}^{-1}$  for optical imaging and  $0''.274 \text{ pixel}^{-1}$  for near-IR imaging. Further information can be found on the ANDICAM instrument specification website.<sup>9</sup>

Our observations of the nova spanned 126 days, starting on 2018 March 25 and ending on 2018 July 29. We took daily cadence from 2018 March 25 to 2018 April 20. After that, we took weekly observations until 2019 July 29. Figure 1 shows a V-band image where we identify the location of the nova and the local field standard. Table 1 presents the optical photometry sequence of the local field. Figure 2 is a J-band image which shows the location of the nova and the field standards. The data



**Figure 1.** V-band image of V906 Car off the Carina Nebula obtained with the CTIO 1.3 m telescope on 2018 April 1. The exposure time was 5 s. The local photometric standard is numbered. The bar corresponds to  $1'$ . North is up, and east is left. The nova is marked in the center of the image and the bright identifiable star HD 92063 north of the nova is labeled.

on the local field standards for the near-IR can be found on Table 2.

We conducted the photometric reduction of the optical data using the Photutils’s aperture photometry tool (Bradley et al. 2016), a part of the Python-based package Astropy (Astropy Collaboration et al. 2013). The use of aperture photometry is appropriate, as the images of the nova are taken with very short exposures and there are no closer sources in the vicinity of the nova.

We estimated the magnitudes for the reference star located in the field of view of the nova using the mean value of five nights of observations. Our photometry for individual nights was computed using the star instrumental magnitude and the *BVR/IR* zero-points derived from observations of photometric standards as described in Wee et al. (2018). We derived the optical zero-points using the Landolt standards in the PG1657+078 field (Landolt 1992). The infrared zero-points are calibrated using 2MASS *JHK* data (Skrutskie et al. 2006). The zero-points are extinction- and color-term corrected based on the data provided on the CTIO’s calibration pages.<sup>10,11</sup>

Due to the brightness of the nova ( $V_{\text{max}} = 5.84 \text{ mag}$ ), only one star is suitable to be used as a reference star in the optical, as it is bright enough to be visible alongside the nova when the object is at its maximum, and dim enough to not be saturated when we increased the exposure time during the later part of the observation. For these reasons, we did not use more reference stars for our optical photometry than the one labeled. We checked the reference star for nonvariability through the data from the Gaia catalog (Gaia Collaboration et al. 2016, 2018), and found that the star has a nominal magnitude error and

<sup>8</sup> <https://www.utias-sfl.net/?p=3015>

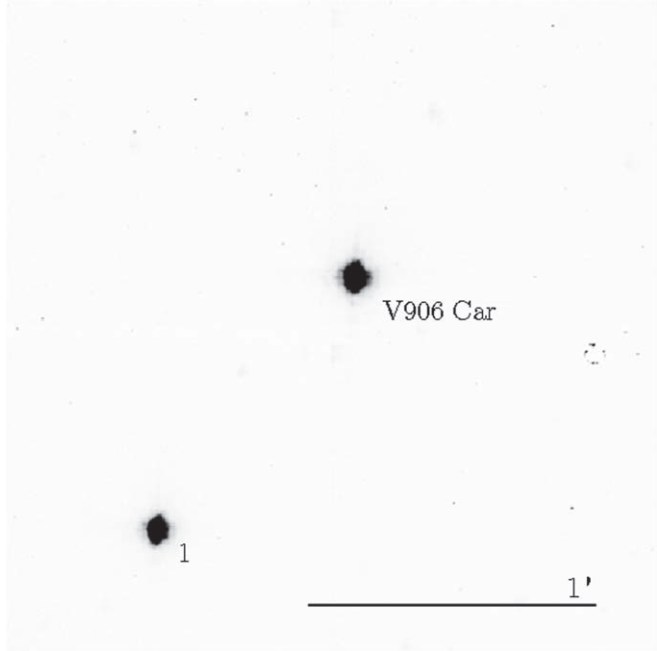
<sup>9</sup> <http://www.astronomy.ohio-state.edu/ANDICAM/detectors.html>

<sup>10</sup> <http://www.ctio.noao.edu/noao/content/13-m-smarts-photometric-calibrations-bvri>

<sup>11</sup> <http://www.ctio.noao.edu/noao/content/photometric-zero-points-color-terms>

**Table 1**  
Optical Photometry Sequence Near V906 Car

| Star ID | $\alpha$ (J2000)                                     | $\delta$ (J2000) | $V$         | $B - V$     | $V - R$     | $V - I$     |
|---------|--|------------------|-------------|-------------|-------------|-------------|
| Nova    | 10 <sup>h</sup> 36 <sup>m</sup> 15 <sup>s</sup> .239 | -59°35'52.326"   | ...         | ...         | ...         | ...         |
| 1       | 10 <sup>h</sup> 36 <sup>m</sup> 22 <sup>s</sup> .371 | -59°36'31.636"   | 9.185 (031) | 1.782 (032) | 1.268 (048) | 2.267 (055) |



**Figure 2.**  $J$ -band image of V906 Car off the Carina Nebula obtained with the CTIO 1.3 m telescope on 2018 April 1. The exposure time was 5 s. The local photometric standards are numbered. The bar corresponds to 1'. The nova is marked in the center of the image. North is up, east is left.

insignificant astrometric noise. There is hence strong evidence that the reference star is not a variable and is appropriate for our photometry calibration purposes.

The near-IR data are reduced using the astronomy software Cyanogen Imaging MaxIm DL's photometric tool, which has the in-built capability to identify, track, and photometer objects across images using aperture photometry. We used the same star in the near-IR  $JHK$  filter magnitudes as we did for the optical images. We present the  $BVRJHK$  light curves in Figure 3.

### 2.2. CTIO 1.5 m CHIRON

We observed V906 Car with the CTIO 1.5 m telescope and CHIRON spectrograph (Schwab et al. 2012; Tokovinin et al. 2013) over 70 days in the “slicer” mode whose spectral resolution is up to  $R = 80,000$ . Upon publication, the data will be made publicly available via the online repository WiseREP (Yaron & Gal-Yam 2012).

### 2.3. Evryscope

We obtained predisccovery  $g'$ -band photometry of the nova from Evryscope, spanning the time between  $T_D - 5$  and  $T_D + 2$ . Evryscope is an array of telescopes capable of observing the entire visible sky down to an airmass of 2 (8150 square degrees), together forming a gigapixel-scale image every two minutes (Law et al. 2015; Ratzloff et al. 2019). Data is typically taken at a 2-min

**Table 2**  
Infrared Photometry Sequence Near V906 Car

| Star ID | $J$         | $H$          | $K$         |
|---------|-------------|--------------|-------------|
| 1       | 5.423 (004) | 4.668 (0135) | 4.304 (011) |

cadence, occasionally twice over if the object appeared in two cameras simultaneously. The images were coadded to a 30-min cadence for better consistency in the PSF. The Evryscope  $g'$  predisccovery light curve is shown in relation to our  $B$  and  $V$  light curves in Figure 4.

Evryscope data is reduced in real time by a custom analysis pipeline described in Ratzloff et al. (2019). After background modeling and subtraction, differential photometry is performed with forced apertures at known source positions in a reference catalog (Ratzloff et al. 2019). The `SysREM` detrending algorithm is applied with two iterations to remove any remaining systematics (Tamuz et al. 2005).

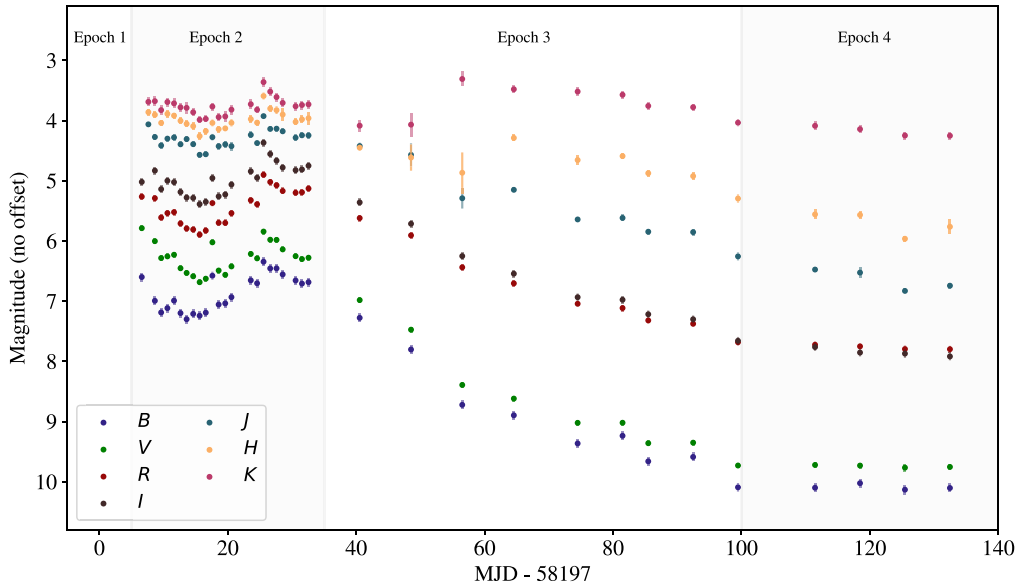
### 2.4. LCO 1.0 m

We obtained a late-time optical photometry of the nova approximately a year after discovery (MJD 58550.2) using the LCO 1.0 m telescope under the program NOAO2019A-011 (PI: N. Blagorodnova). The data were reduced using the LCO BANZAI pipeline (McCully et al. 2018). Aperture photometry on the target was performed using `Python photutils` routines. The zero-point for  $g$  and  $r$  bands were calibrated using isolated field stars that were available in the SkyMapper DR1.1 catalog (Wolf et al. 2018). The zero-points for  $B$  and  $V$  bands were derived using field stars present in the UCAC4 catalog (Zacharias et al. 2012). The rest of filters was calibrated using as a reference the star from Table 1.

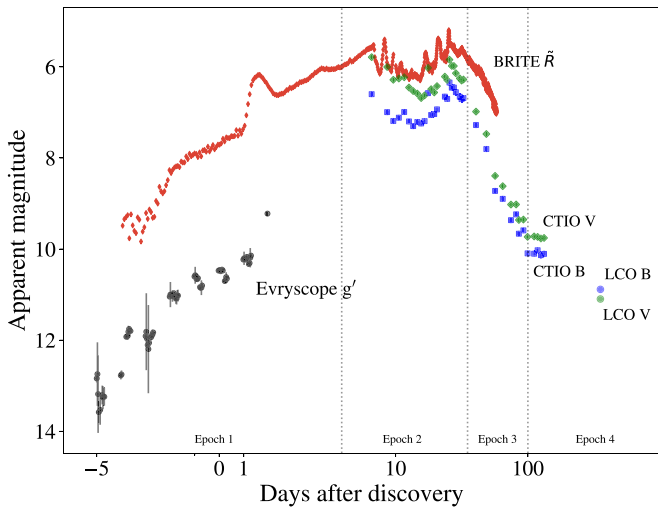
## 3. Extinction Estimate

Using the Milky Way extinction law from Fitzpatrick (1999) and the extinction parameters from Schlafly & Finkbeiner (2011), we obtain an  $A_V = 3.598$  for the position of the nova. However, this shall be only considered as an upper limit, as this is the total line-of-sight extinction in our own Galaxy; the nova, located approximately at the distance of the Carina Nebula, has likely less extinction. To obtain a better estimate for the foreground extinction to our target, we used the strength of Na I lines from high-resolution spectroscopy of the nova as a proxy for the extinction.

We used the high-resolution spectrum taken with the CTIO 1.5 m telescope to measure the absorption in the wavelengths of Na I  $D_1$  and Na I  $D_2$  using the IRAF spectral analysis routines. The absorption was assumed to arise from foreground interstellar medium. After measuring the equivalent widths, we computed the average values for both the  $D_1$  and  $D_2$  lines. Our sample of Na I equivalent width measurements was converted into  $E(B - V)$  using the empirical relations found in Poznanski et al. (2012). We



**Figure 3.** Light curves of V906 Car in the optical and near-infrared in their constituent epochs. No offset is applied. (The data used to create this figure are available.)



**Figure 4.** Prediscovery early-time data from Evryscope. BRITE  $\tilde{R}$  data is superimposed on CTIO  $B$  and  $V$  photometry. Late-time LCO  $B$  and  $V$  data are also added. No offsets are applied.

(The data used to create this figure are available.)

also calculated  $A_V$  using the assumed ratio for total to selective extinction  $R_V$  of 3.1 (Cardelli et al. 1989). The result of our analysis suggests a foreground extinction toward V906 Car of  $A_V = 1.116^{+0.538}_{-0.385}$  mag, which is likely unaffected by the peculiar reddening law inside the nebula (e.g.,  $R_V = 4.8$ ; Smith 1987). Our value is in good agreement with the foreground extinction values toward Carina derived in recent studies (Hur et al. 2012), which estimated a relatively low reddening  $E(B - V) \sim 0.36 \pm 0.04$ , and extinction of  $A_V \sim 1.1$  mag. These values also agree with the spectroscopic analysis done by Aydi et al. (2020), which used combined Na I  $D_1$  and  $D_2$  interstellar absorption doublet method and DIBs to derive a foreground extinction toward the nova of  $A_V = 1.11 \pm 0.05$  mag.

## 4. Photometric Evolution of V906 Car

### 4.1. Evolution of Optical and Near-IR Emission

The nova conforms to the general morphology of a nova light curve as described by Bode & Evans (2008). The nova is observed to have an initial rise from  $t_{D-5}$  to  $t_{D+7}$ , a premax halt at  $t_{D+6}$ , followed by a final maximum on  $t_{D+26}$ , an early decline from  $t_{D+26}$  to  $t_{D+100}$ , and finally a slow down in the decline after  $t_{D+100}$ .

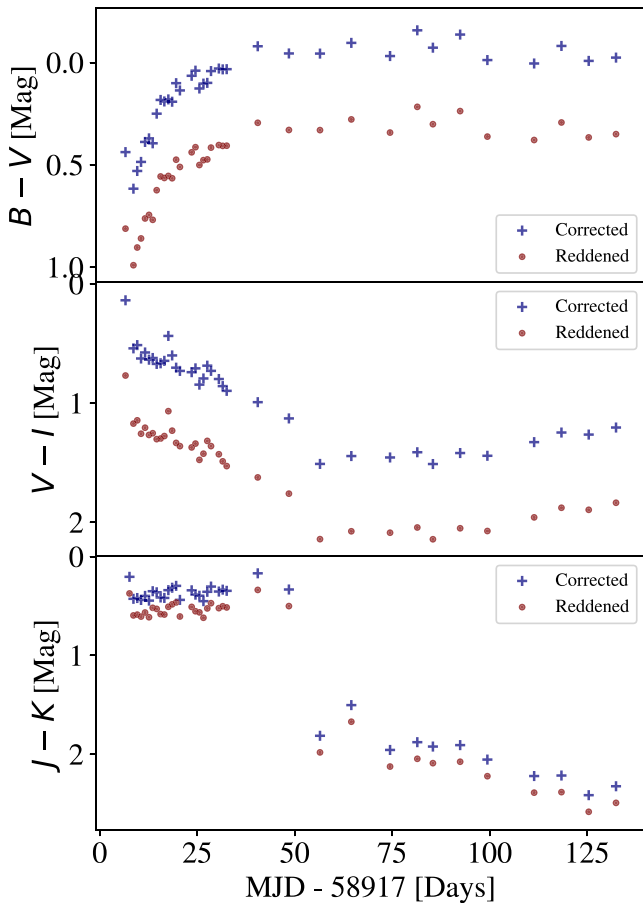
The photometric evolution of V906 Car can be similarly grouped into four epochs: (1) the rising part from  $t_{D-5}$  to  $t_{D+2}$ ; (2) the evolution at peak, lasting from  $t_{D+5}$  to  $t_{D+35}$ ; (3) the decay phase from  $t_{D+35}$  to  $t_{D+100}$ ; and (4) the late plateau phase, which begins at  $t_{D+100}$ .

In the first epoch, the Evryscope data (Figure 4) shows that the nova rose from  $g' = 13.58 \pm 0.23$  to  $g' = 9.22 \pm 0.03$  mag in 6 days. The light curve shows minor bumps of fractions of a magnitude along a power-law rise. The nonsmooth uniform rise seems to qualitatively agree with the prepeak high cadence data obtained by BRITE.

In the second epoch, the light curve begins with an immediate decline from our first day of observations. The decline lasts for 10 days in all filters, falling most significantly in the  $B$  and  $V$  with a  $\sim 1$  mag decline. The dip, however, was not significant in the  $JHK$  filters. After this initial decline, the nova rebrightened to its peak in all  $BVRJHK$  filters, with  $V_{\max} = 5.84 \pm 0.09$  mag. Another interesting feature is that the early period is punctuated with sudden outbursts, with the most pronounced flare on  $t_{+15}$ , where the nova quickly brightened for about a day before falling back to a slower rise in all filters. The flare was most significant in the optical filters, brightening by 0.6 mag in the  $V$  band.

In the third epoch, the brightness of the nova starts to decline at  $t_{+35}$ . The decline lasts until  $t_{100}$  for all bandpass filters except for the  $H$  and  $K$  bands, which begin to brighten at around  $t_{50}$ , before reaching a peak and declining thereafter at  $t_{60}$ . Notably, the  $K$  magnitude reached a maximum brighter than the initial maximum in the early epoch. This change is probably unrelated





**Figure 5.** Color evolution  $B - V$ ,  $V - I$ , and  $J - K$  of V906 Car. Extinction corrections are based on the extinction value of  $A_V = 1.11$ .

to dust formation, as we do not see a significant rise in the other near-IR bands. A more likely explanation is an increase in the flux in the He I  $2.05 \mu\text{m}$  line emission (Shore et al. 1996). The larger excess in  $K$  is consistent with expectations that the thermal emission will peak in or beyond the  $K$  band.

In the fourth and final epoch, the brightness of the nova remains largely flat in the  $BVRI$  filters until the end of our observations. The brightness of the nova continue to decline in the  $JHK$  bands, with  $J$  and  $H$  declining at a similar rate, while  $K$  declines at a much slower rate.

#### 4.2. Color Evolution

We present the evolution of the nova optical and infrared colors  $B - V$ ,  $V - I$ , and  $J - K$  in Figure 5. In the  $B - V$  curve, the nova becomes bluer from day 10 to day 45 after outburst, and then levels off to a nearly constant value of  $B - V = 0.24$  for the remaining observations. The bluer colors are accompanied by a steady reddening in  $V - I$ , which changes from  $V - I = 1.3$  to  $V - I = 2.3$  during the same period. The infrared color index  $J - K$  shows a quick change near day 50, and shifts from steady values of  $J - K = 0.7$  observed prior to day 50 to values above  $J - K = 2$  after day 50. At later times, while the optical  $B - V$  color remains stable and the  $V - I$  color gets bluer, the IR bands slowly evolve toward redder colors.

We interpret these changes in optical and infrared color in terms of the thinning of the photospheric optical depth, which would result in an unobstructed view of the hotter accretion disk and the WD, having bluer  $B - V$  colors. The later

**Table 3**  
Photometry of the Progenitor for V906 Car

| Survey | Band      | Magnitude (mag)  | Reference |
|--------|-----------|------------------|-----------|
| VPHAS+ | $u$       | $19.46 \pm 0.05$ | (1)       |
| VPHAS+ | $g$       | $20.13 \pm 0.03$ | (1)       |
| VPHAS+ | $H\alpha$ | $19.09 \pm 0.05$ | (1)       |
| VPHAS+ | $r$       | $19.62 \pm 0.03$ | (1)       |
| VPHAS+ | $i$       | $19.18 \pm 0.04$ | (1)       |
| VISTA  | $J$       | $18.58 \pm 0.06$ | (2)       |
| VISTA  | $H$       | $18.41 \pm 0.10$ | (2)       |

**Notes.** All magnitudes are reported in the vega system.

<sup>a</sup> References: (1) Drew et al. (2014, 2016), (2) Preibisch et al. (2014).

production of dust would cause a gradual increase in  $V - I$  colors from increasing infrared emission from heated dust. The sudden increase in  $J - K$  colors at day 45 could be explained by the formation of this dust envelope within the nova, which over time causes an increasing infrared excess from the source.

## 5. Analysis

The progenitor of V906 Car lies at a closer angular separation from the Carina Nebula, suggesting a possible association. The distance to the nebula was estimated to be of  $d = 2.3 \pm 0.5$  kpc using the proper motion of the Homunculus Nebula (Smith 2006). The distance for the progenitor star as provided in the Gaia DR2 distance catalog (Bailer-Jones et al. 2018; Gaia Collaboration et al. 2018) is  $d = 3.35^{+2.81}_{-1.60}$  kpc. Although the errors are large due to the faint magnitude of the star, this value seems to be consistent with being associated with the nebula. In fact, our estimate for the extinction is significantly lower than the total value for the line of sight. Because the nebula is likely to contribute significantly to the overall reddening value, in our analysis we will assume that the nova is most likely located in front of or inside the nebula, and we adopt the nebula’s distance for the progenitor.

### 5.1. Progenitor Modeling

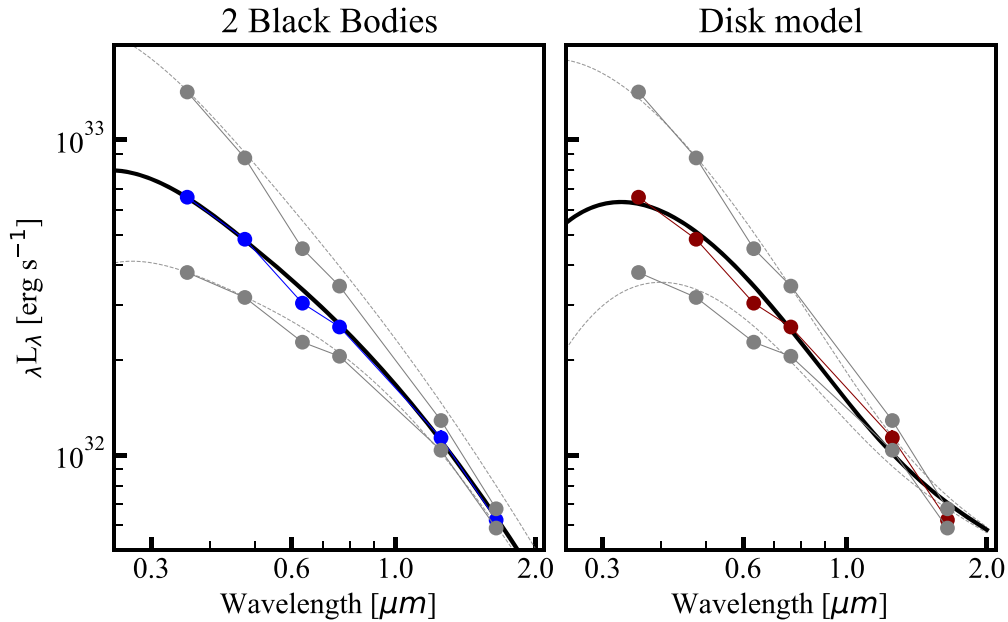
The location of the nova was covered by the VST Photometric  $H\alpha$  Survey of the Southern Galactic Plane and Bulge (VPHAS+; Drew et al. 2014), performed between 2012 February and 2013 April (7–8 yr before the nova outburst) and the Visible and Infrared Survey Telescope for Astronomy (VISTA; Preibisch et al. 2014), which observed the area in 2012 March.

We used the online portal *Vizier*<sup>12</sup> to retrieve the archival magnitudes of the progenitor system, which are shown in Table 3.

The modeling of the progenitor was performed using a custom developed code *BBFit*,<sup>13</sup> based on Markov-Chain Monte Carlo *Python* implementation in the package *emcee* (Foreman-Mackey et al. 2013). After correcting for the foreground extinction (see Section 3), we used the transmission curves for each filter—available through the Spanish VO Filter Profile Service (Rodrigo et al. 2012)—to obtain the monochromatic flux for each band. Initially, we modeled the emission with two blackbody components. Due to an excess

<sup>12</sup> <http://vizier.u-strasbg.fr/viz-bin/VizieR>

<sup>13</sup> <https://github.com/nblago/utlis>



**Figure 6.** Best fit to the progenitor spectral energy distribution (SED) with a two-blackbody model (left) and a steady accretion disk model (right). The filled points correspond to the archival flux for the progenitor in the *ugrizH* bands. The black line corresponds to the best-fit model for  $A_V = 1.116$ . The gray markers show the SED with extinction correction values of  $A_V = 0.73$  and  $A_V = 1.65$  mag. The dashed gray lines correspond to the best-fit models for those SEDs.

**Table 4**  
Best-fit Parameters for the Progenitor Spectral Energy Distribution (SED) Modeling

| $A_V$<br>(mag) | Two Blackbodies         |                                  |   |                         |                                  |   | Disk Model             |                       |  |                                   |  |
|----------------|-------------------------|----------------------------------|---|-------------------------|----------------------------------|---|------------------------|-----------------------|--|-----------------------------------|--|
|                | $T_{\text{BB1}}$<br>(K) | $R_{\text{BB1}}$<br>( $10^8$ cm) | $L_{\text{BB1}}$<br>( $10^{-1} L_\odot$ ) | $T_{\text{BB2}}$<br>(K) | $R_{\text{BB2}}$<br>( $10^8$ cm) | $L_{\text{BB2}}$<br>( $10^{-2} L_\odot$ ) | $M_*$<br>( $M_\odot$ ) | $R_*$<br>( $10^8$ cm) | $\log$<br>[ $\dot{M}/(M_\odot \text{yr}^{-1})$ ] | $R_{\text{out}}$<br>( $R_\odot$ ) |  |
| 0.731          | $13936^{+2188}_{-1395}$ | $45^{+9}_{-9}$                   | $1.4^{+2.3}_{-0.8}$                       | $5067^{+319}_{-351}$    | $179^{+9}_{-7}$                  | $4.0^{+1.6}_{-1.2}$                       | $0.72^{+0.23}_{-0.23}$ | $27.2^{+0.5}_{-0.5}$  | $-9.7^{+0.2}_{-0.1}$                             | $20.7^{+0.6}_{-0.6}$              |  |
| 1.116          | $15073^{+2124}_{-1450}$ | $54^{+9}_{-9}$                   | $2.8^{+3.8}_{-1.6}$                       | $4951^{+497}_{-527}$    | $172^{+14}_{-11}$                | $3.3^{+2.4}_{-1.5}$                       | $0.71^{+0.23}_{-0.19}$ | $21.2^{+0.5}_{-0.5}$  | $-9.7^{+0.2}_{-0.1}$                             | $22.1^{+0.7}_{-0.6}$              |  |
| 1.654          | $18481^{+1502}_{-1040}$ | $61^{+4}_{-5}$                   | $8.4^{+4.7}_{-2.9}$                       | $3904^{+616}_{-551}$    | $192^{+38}_{-25}$                | $1.6^{+2.5}_{-0.9}$                       | $0.73^{+0.22}_{-0.22}$ | $13.1^{+0.4}_{-0.4}$  | $-9.7^{+0.2}_{-0.1}$                             | $24.9^{+0.7}_{-0.7}$              |  |

of flux in the  $H\alpha$  band, likely due to an existing emission line, we left this band out of the analysis. The best-fit model for different values of extinction is depicted in Figure 6 and the posterior parameters are summarized in Table 4.

Our results show that the hot component has a temperature of  $\sim 1.5 \times 10^4$  K, and a luminosity of  $\sim 3 \times 10^{-1} L_\odot$ . Its estimated radius of  $\sim 50 \times 10^8$  cm is larger than the radius of a single WD, which ranges  $0.5 \times 10^8 \leq R_{\text{WD}} \leq 14 \times 10^8$  cm. These values are in agreement with a bright spot or an accretion disk, rather than the WD itself.

The cold component has a luminosity of  $(1.6\text{--}4) \times 10^{-2} L_\odot$ , consistent with a cold dwarf classification. However, its estimated temperature is higher than the one expected for a main-sequence star within this luminosity range. Most likely, this is due to a reflection effect, when in a tidally locked system the radiation from the WD heats one of the sides of the cold companion.

Given the donor is a main-sequence star, we use the mass-to-radius and the mass-to-luminosity relations to estimate its mass. For our canonical extinction value, the  $R/R_\odot = (M/M_\odot)^{0.8}$  relation suggests a donor with  $\simeq 0.25 M_\odot$ . When using the mass-luminosity relation  $L/L_\odot \simeq 0.23(M/M_\odot)^{2.3}$ , we derive an upper limit for the companion mass of  $\simeq 0.43 M_\odot$ .

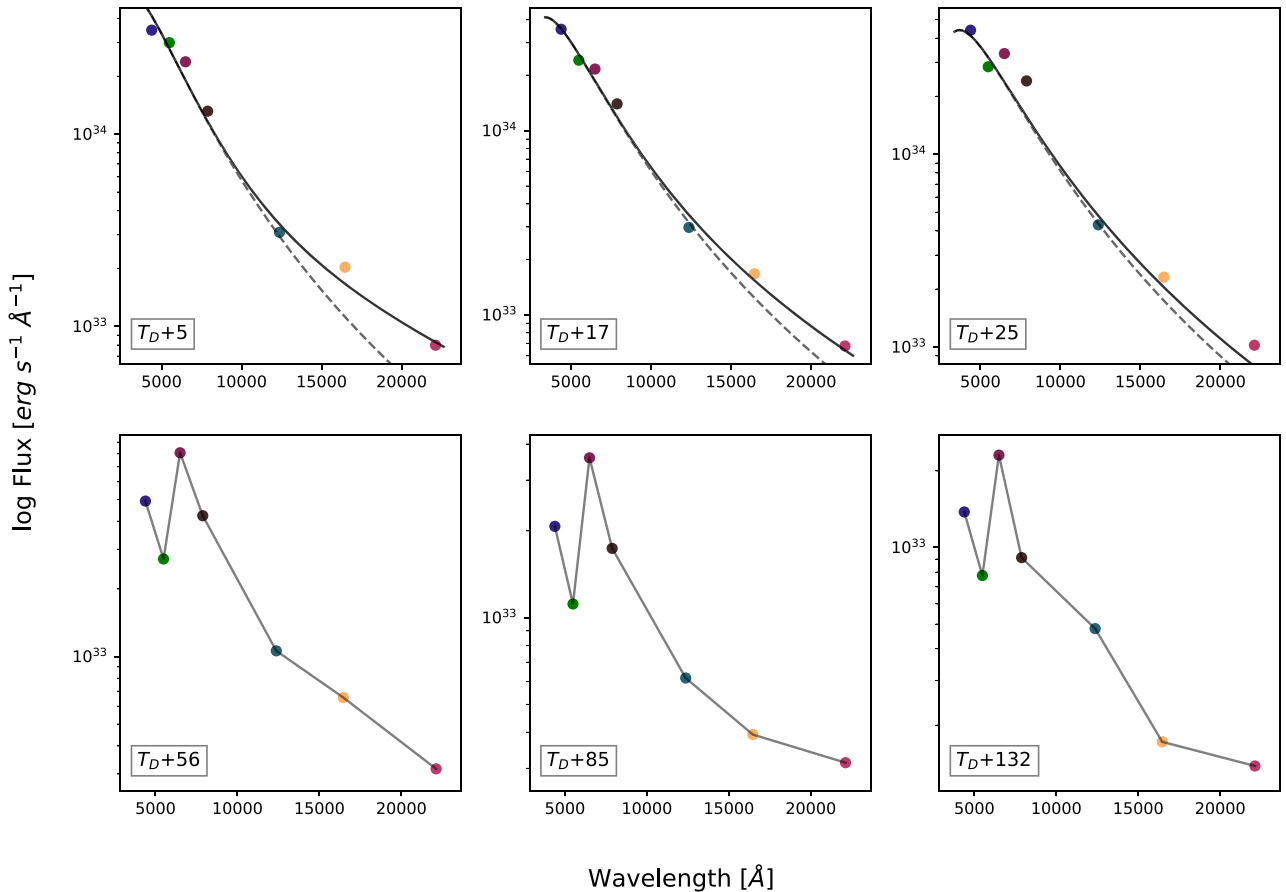
Provided the nova progenitor was an accreting WD system, we undertake a second analysis of the observed SED using the formalism of a viscous accretion disk model developed by Shakura & Sunyaev (1973) and Lynden-Bell & Pringle (1974).

In this approach, we assume that the emission is dominated by the accretion disk and therefore the integrated SED of the source is composed by several concentric optically thick annuli, each one radiating at a different temperature. Assuming that the accretion rate through the disk is steady, the energy dissipated within each annulus is independent of the viscosity. For a specific implementation, we adopt the model described in Frank et al. (2002), where the temperature of each annulus at radius  $R$  corresponds to

$$T(R) = \left\{ \frac{3GM\dot{M}}{8\pi R^3 \sigma} \left[ 1 - \left( \frac{R_*}{R} \right)^{1/2} \right] \right\}^{1/4}, \quad (1)$$

where  $M_*$  is the mass of the star,  $R_*$  the stellar radius,  $\dot{M}$  the accretion rate, and  $\sigma$  and  $G$  are the Stefan-Boltzmann and gravitational constants. Following the approach of Kenyon et al. (1988), which used the steady disk approach to study FU Orionis, we assume that the disk emits at its maximum temperature for any radius  $R_* \leq R \leq 1.5R_*$ .

The spectrum for each annulus is approximated as a blackbody at effective temperature  $l_\nu = B_\nu[T(R)]$ , which is integrated up to the outer radius of the disk  $R_{\text{out}}$ . The binary inclination angle is designed by  $i$ , which we set to a fix value of  $i = 45^\circ$ . The emission from each ring of thickness  $dR$  is scaled by the solid angle as seen from distance  $D$ ,  $2\pi R dR \cos i/D^2$ . This provides



**Figure 7.** Evolution of the SED of the nova. The fluxes at different epochs after discovery are shown for *BVRJIHK* bands. The dotted lines represent single blackbody fit, while the solid lines represent best fit for two-blackbody model. After  $t_{50}$ , the SED is dominated by strong emission lines, which cause the SED to deviate from a blackbody.

the total monochromatic flux to be

$$F_\nu = \frac{4\pi h \cos i \nu^3}{c^2 D^2} \int_{R_*}^{R_{\text{out}}} \frac{R dR}{e^{h\nu/kT(R)} - 1}. \quad (2)$$

This accretion disk model has four different parameters, for which we assume uniform priors:  $0 \leq M_* \leq 1.44 M_\odot$ ,  $0 \leq R_* \leq 10 R_\odot$ ,  $-10 \leq \log[\dot{M}/(M_\odot \text{ yr}^{-1})] \leq -5$  and  $R_* \leq R_{\text{out}} \leq 50 R_\odot$ .

Adopting different foreground extinction values for the system, the best-fit parameters of the disk model are shown in Table 4. For our best estimate of the extinction  $A_V = 1.116$ , the computed stellar radius is still a factor of two larger than the one expected for WDs of the given mass. However, deviations are not unexpected, due to the highly simplified disk model used in our analysis and ignoring the contribution of the cold donor star. Also, assuming a larger value for the extinction  $A_V$  also allows us to obtain a more consistent radius, while all the other parameters are not significantly altered.

The best-fit SEDs for both models are shown in Figure 6. The two-blackbody model seems to already work reasonably well in predicting both the blue flux from the disk and the colder contribution from the companion. The composite SED of the accretion disk seems to overpredict the flux at longer wavelengths, which translates in a larger  $R_{\text{out}}$  than the one physically expected for such system. The progenitor mass that we obtain with the accretion disk modeling does not show a

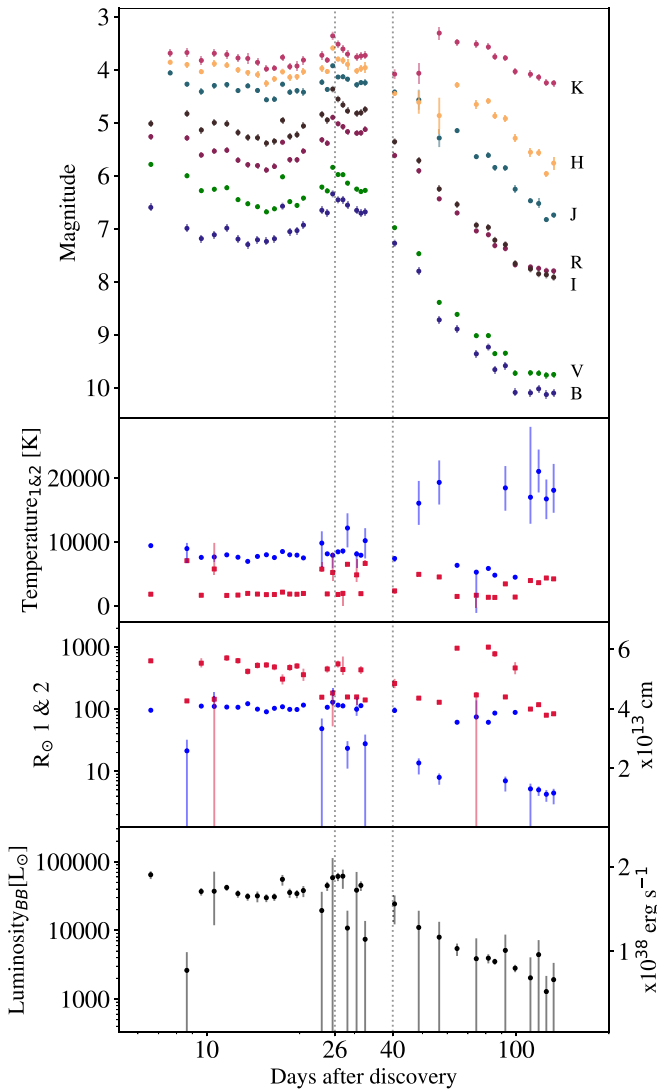
strong dependence on the extinction. This estimate is also in agreement with the WD mass derived from our analysis of the decay of the light curve, as further explained in Section 6.2.

### 5.2. Evolution of the V906 Car Eruption

Using the same assumptions for distance and extinction as in Section 5.1, we used *BBFit* to model the nova SED at different phases, shown in Figure 7. As mentioned earlier, because soft X-ray emission was not detected from the nova, we worked only with optical and near-IR data and did not include X-ray or UV data in our modeling. Provided the existence of an IR excess component in nova eruptions, we chose to use two blackbodies to model the emission.

The best-fit measurements from our analysis are shown in Figure 8, along with the multiband photometry. We divide our analysis into three different phases based on the evolution of the SED. In the first phase, from  $t_D$  to  $t_{D+25}$ , the nova is dominated by the optically thick wind, which agrees fairly well with a blackbody description.

The temperature and radii derived using this simple model remain fairly consistent. The average temperature of the hotter blackbody is  $\sim 9540$  K and its photospheric radius is  $\sim 70R_\odot$ , while the colder is at 2440 K throughout the first 25 days (epoch2), with a radius of  $\sim 500R_\odot$ . The luminosity of the system is also fairly consistent in these 25 days, with a median of  $3.6 \times 10^4 L_\odot$  (or average of  $3.96 \times 10^4 L_\odot$ ). Since the peak of the emission is at shorter wavelengths, without UV data, it is

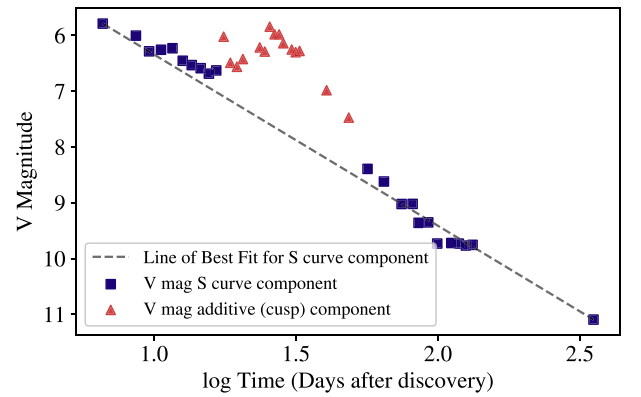


**Figure 8.** Temperature, radii, and luminosity of two-blackbody fit modeling over time. The hotter component is shown with blue markers, and the colder components with green. The dotted line indicates the transition of the different phases in the nova SED evolution.

difficult to exactly constrain the SED. Therefore, our estimate of the nova’s bolometric luminosity should be interpreted as a lower limit without including the emission at shorter wavelengths.

In the second phase of the SED evolution, from  $t_D + 25$  to  $t_D + 40$ , the source starts to develop an optically thin wind with strong emission lines, specially in the *R*- and *I*-band regions (see Figure 6), which makes the blackbody fits less reliable. In our attempt to minimize the effect of the lines, we decide to exclude these bands from our fit. However, the larger scatter and increased error bars on the best-fit parameters indicate that these values should be interpreted with caution.

In the third and fourth phases, from  $t_D + 40$  onward, the nonblackbody-like appearance of the nova SED makes the interpretation challenging, as shown in Figure 7. The emission is not well approximated by a simple blackbody model. At this epoch, the hotter component increases in temperature to  $\sim 2 \times 10^4$  K, as the photosphere recedes closer to the WD. The bolometric luminosity shows a smooth decline. While the fits qualitatively agree with the nova expected behavior, a more careful future analysis based on spectroscopic modeling is required.



**Figure 9.** Underlying smooth nova characteristics fitted with trend line, emphasizing additive component.

## 6. Discussion

### 6.1. Nova Classification

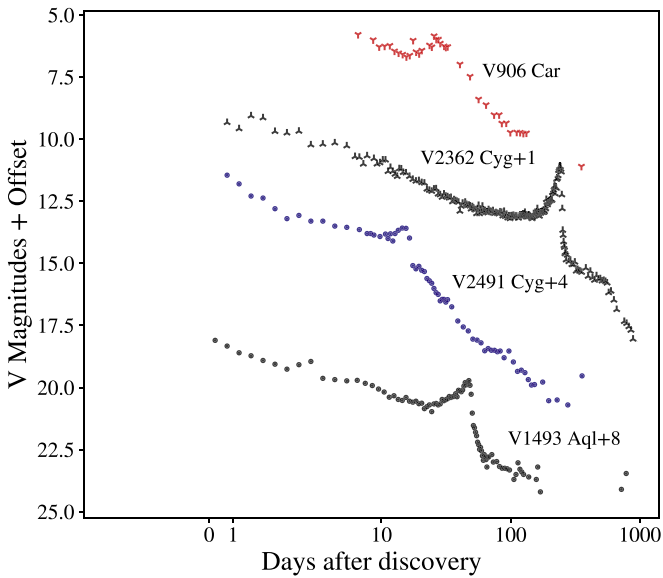
From the evolution of V906 Car in the visual band, we derive the times for the nova to fade by 2 and 3 mag,  $t_{2,V} = 26.2$  days and  $t_{3,V} = 33.0$  days. The decline rate (speed) makes V906 Car a “fast nova” according to the Gaposchkin (1957) classification. Categorizing CNe in speed classes is, however, for most parts only useful taxonomically. A nova’s astrophysical significance can be better understood in relation to the classification method by Stroepe et al. (2010), which categorizes CNe according to their light-curve morphology.

V906 Car exhibits characteristics of various nova classification in terms of light-curve shape. An early photometric analysis by Domingo et al. (2018) determined V906 Car to be a J class nova, due to the “jittering” above the base level of the nova light-curve shape. Domingo et al. (2018) also considered the potential that V906 Car could be classified as a O class nova, but commented that the oscillations in a O class nova are supposed to be quasi-periodic, and the periodicity of the jitters in V906 Car appeared to be random.

The analysis by Domingo et al. (2018) is based on a limited profile of the nova, with less than 3 months of data. Using our combined CTIO and LCO data spanning for about a year, we consider V906 Car to most closely resemble a C class nova instead. C class novae display characteristic cusps in their light curves, which appear as secondary additive on top of the base level S class type light-curve shape (Stroepe et al. 2010). Figure 9 shows how there appears to be a strong additive component atop the expected S class curvature for V906 Car. Approximately 30 days after discovery, the nova reaches a peak rapidly at  $m_V = 5.84 \pm 0.09$  before resuming to the decline law of S class light-curve shape.

Our calculations of the decline rates (in units of magnitudes per logarithmic time) of V906 Car post primary peak yields the following results for each of its declining phases (see Table 5) show that its decline characteristics conform to declines observed in C class novae. Our calculations of the strength of the cusps (see Table 6) also show that V906 Car conforms to the morphology of a C class nova, with an increase from  $V_{\min}$  to  $V_{\text{cusp}}$  of 0.78 mag and a significant decrease from  $V_{\text{cusp}}$  to  $V_{\text{base}}$  of 2.55 mag. Comparatively, V906 Car resembles V2491 Car the most, with similar decline patterns and smaller cusps relative to their general decline. V906 Car’s smaller cusp compared to V2352 Cyg and V1492 Aql is consistent with the





**Figure 10.** Comparison of ASASSN-18fv with other C class novae in logarithmic time with LCO 1.0 m V.

**Table 5**  
Slope Comparison with Other C Class Novae

| Nova                   | Slope 1 | Slope 2 | Slope 3 |
|------------------------|---------|---------|---------|
| V1493 Aql <sup>a</sup> | -4.0    | -2.0    | ...     |
| V2362 Cyg <sup>a</sup> | -2.7    | -2.7    | -10.8   |
| V2491 Cyg <sup>a</sup> | -2.0    | -7      | -3.8    |
| V906 Car               | -2.24   | -3.31   | -3.94   |

**Note.**

<sup>a</sup> Figures presented are from Strope et al. (2010).

**Table 6**  
Cusp Comparison with other C class Novae in V Magnitude

| Nova                   | $V_{\min}$ | $V_{\text{cusp}}$ | $V_{\text{base}}$ |
|------------------------|------------|-------------------|-------------------|
| V1493 Aql <sup>a</sup> | 12.8       | 11.7              | 14.9              |
| V2362 Cyg <sup>a</sup> | 12.1       | 10.1              | 13.6              |
| V2491 Cyg <sup>a</sup> | 9.9        | 9.6               | 11.2              |
| V906 Car               | 6.6        | 5.8               | 8.39              |

**Note.**

<sup>a</sup> Figures presented are from Strope et al. (2010).

expectations by Strope et al. (2010) that the later the cusp appears, the stronger it is. A comparison of V906 Car in V band with other C class novae is presented in in Figure 10.

### 6.2. WD Mass

Optical and infrared light curves of CNe are approximately homologous among various WD masses and chemical compositions when free-free emission from optically thin ejecta is spherical and dominates the continuum flux of novae (Hachisu & Kato 2006). Such a homologous template light curve is called “a universal decline law” (Hachisu & Kato 2006). The timescale of the light curve depends strongly on the WD mass but weakly on the chemical composition, and hence we are able to roughly estimate the WD mass from the light-curve fitting.

The template light curve for the universal law has a slope of flux  $F \propto t_{\text{init}}^{-1.75}$  in the middle from  $t_2$  to  $t_6$  below optical maximum, but it declines more steeply where  $F \propto t_{\text{init}}^{-3.5}$  from  $t_6$  to  $t_{10}$  (Hachisu & Kato 2006). The point of which the decline rate changes, caused by a quick decrease in the wind mass-loss rate, is known as the break time  $t_{\text{break}}$  caused by a quick decrease in the wind mass-loss rate (Hachisu & Kato 2006). Once  $t_{\text{break}}$  is determined, we can derive the period of a UV burst phase, the duration of the optically thick wind phase, and the turnoff date of hydrogen shell-burning (Hachisu & Kato 2006).

As noted in Aydi et al. (2020), the X-ray emission from V906 Car has been heavily attenuated and reprocessed into optical and infrared light, both from absorption and X-ray suppression in corrugated shock fronts (Metzger et al. 2015; Steinberg & Metzger 2018). This makes it difficult to use the X-ray emission to constrain the WD properties, and available X-ray data reported in Aydi et al. (2020) included a non-detection in the 0.3–10.0 keV range with a  $3\sigma$  upper limit. To provide some constraints on the WD mass, we compared the late-time optical emission with the grid of models in Hachisu & Kato (2006), recognizing the limitations and uncertainties inherent in using only optical data and possible systematic differences between the V906 Car nova from the novae used in calibrating the grid of models within Hachisu & Kato (2006).

From studying our observed optical photometry, we estimate that  $\log t_{\text{break}} = 2.67$ , or  $\sim 470$  days. Using a grid of models of CO novae from the literature, we observe that our break time provides an upper limit to the WD mass of  $M_{\text{WD}} < 0.8M_{\odot}$ , for compositions with  $0.35 < X < 0.70$ , and CNO metallicity between  $0.03 < X_{\text{CNO}} < 0.3$  (Hachisu & Kato 2006).

## 7. Conclusion

We have presented optical and near-IR photometry of V906 Car. We found that the nova is a fast nova, peaking at  $V_{\text{max}} = 5.84 \pm 0.09$  mag and has rapid decline times of  $t_{2,v} = 26.2$  days and  $t_{3,v} = 33.0$  days. We have also presented pre-discovery data from Evryscope, showing that the nova is already visible at  $g' \simeq 13$  mag five days before discovery, and the early photometry provides rare data of evolution of this event well before the peak magnitude is observed.

Our spectroscopic analysis yields a reddening of  $E(B - V) \sim 0.36 \pm 0.04$  and an extinction of  $A_V \sim 1.116^{+0.538}_{-0.385}$  mag in the line of sight to the nova, making it a likely member of the Carina Nebula.

Our modeling of the nova progenitor in quiescence shows a good agreement with the emission coming from two blackbodies. The hotter component has  $T_{\text{BB1}} \sim 1.5 \times 10^4$  K and  $L_{\text{BB1}} \simeq 2.8 \times 10^{-1} L_{\odot}$ . With a radius of  $R_{\text{BB1}} \sim 50 \times 10^8$  cm, this emission is likely associated with the accretion disk. The colder component represents the donor star. Its luminosity  $L_{\text{BB2}} \sim 3 \times 10^{-2} L_{\odot}$  and radius  $R_{\text{BB2}} \sim 1.7 \times 10^{10}$  cm suggest that the star is a  $0.23\text{--}0.42 M_{\odot}$  K-M type dwarf which is likely being heated by the WD companion. When a steady accretion disk model is used to interpret the system in quiescence, the analysis yields a progenitor mass of  $M_{*} = 0.71^{+0.23}_{-0.19} M_{\odot}$  and radius of  $R_{*} = 21.2^{+0.5}_{-0.5} \times 10^8$  cm, consistent with a WD.

An unambiguous classification for the nova is difficult, but overall the nova most closely resembles a C class nova, with a steep decline slope of  $\alpha = -3.94$ . The steep decline slope is consistent with other C class novae, such as V2362 Cyg, V2491 Cyg, and V1493 Aql. Our data also allows us to provide

an additional constraint for the upper limit on the WD mass of  $M_{\text{WD}} < 0.8M_{\odot}$ .

We have used combined infrared and optical data spanning a wide range of times from well before peak magnitude to months afterwards to characterize the nearby V906 Car nova and its progenitor system. Future work using complementary data in other wavelengths will help further refine the characteristics of this peculiar transient and help to improve current astrophysical models for these systems.

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*Facilities:* CTIO:1.3 m, 1.5 m, Evryscope, LCO: 1 m.

*Software:* Astropy (Astropy Collaboration et al. 2013), Cyanogen Imaging MaxIm DL v6 (<http://diffractionlimited.com/product/maxim-dl/>), LCO BANZAI pipeline McCully et al. (2018), BBFit (<https://github.com/nblago/utills>), emcee (Foreman-Mackey et al. 2013).

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## References

- Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, **558**, A33
- Aydi, E., Sokolovsky, K. V., Chomiuk, L., et al. 2020, *NatAs*, **4**, 776
- Bailer-Jones, C. A. L., Rybizki, J., & Fouesneau, M. 2018, *Mantelet*, G., Andrae, R, *AJ*, **156**, 58
- Bode, M. F., & Evans, A. 2008, *Classical Novae*, Vol. 3 (2nd ed.; Cambridge: Cambridge Univ. Press)
- Bradley, L., Sipocz, B., Robitaille, T., et al. 2016, *Photutils: Photometry Tools v0.6*, Astrophysics Source Code Library, ascl:1609.011
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, **345**, 245
- Domingo, A., Hernanz, M., Kuulkers, E., & Kuin, P. 2018, *ATel*, **11677**, 1
- Drew, J. E., Gonzalez-Solares, E., Greimel, R., et al. 2014, *MNRAS*, **440**, 2036
- Drew, J. E., Gonzalez-Solares, E., Greimel, R., et al. 2016, *yCat*, **2341**, 0
- Fitzpatrick, E. L. 1999, *PASP*, **111**, 63
- Foreman-Mackey, D., Hogg, D. W., Lang, D., & Goodman, J. 2013, *PASP*, **125**, 306
- Frank, J., King, A., & Raine, D. J. 2002, *Accretion Power in Astrophysics: Third Edition* (Cambridge: Cambridge Univ. Press), 398
- Gaia Collaboration, Brown, A. G. A., Vallenari, A., et al. 2018, *A&A*, **616**, A1
- Gaia Collaboration, Prusti, T., de Bruijne, J. H. J., et al. 2016, *A&A*, **595**, A1
- Gaposchkin, C. H. P. 1957, *The Galactic Novae* (New York: Interscience)
- Hachisu, I., & Kato, M. 2006, *ApJS*, **167**, 59
- Hounsell, R., Darnley, M. J., Bode, M. F., et al. 2016, *ApJ*, **820**, 104
- Hur, H., Sung, H., & Bessell, M. S. 2012, *AJ*, **143**, 41
- Izzo, L., Molaro, P., Mason, E., et al. 2018, *ATel*, **11468**, 1
- Kenyon, S. J., Hartmann, L., & Hewett, R. 1988, *ApJ*, **325**, 231
- Landolt, A. U. 1992, *AJ*, **104**, 340
- Law, N. M., Fors, O., Ratzloff, J., et al. 2016, *Proc. SPIE*, **9906**, 99061M
- Law, N. M., Fors, O., Ratzloff, J., et al. 2015, *PASP*, **127**, 234
- Liller, W., Shao, C. Y., Mayer, B., et al. 1975, *IAUC*, **2848**, 4
- Lucas, P. 2018, *ATel*, **11460**, 1
- Lynden-Bell, D., & Pringle, J. E. 1974, *MNRAS*, **168**, 603
- McCully, C., Turner, M., Volgenau, N., et al. 2018, *Lcogt/Banzai: Initial Release, 0.9.4*, Zenodo, doi:10.5281/zenodo.1257560
- Metzger, B. D., Finzell, T., Vurm, I., et al. 2015, *MNRAS*, **450**, 2739
- Nelson, T., Mukai, K., Sokolowski, J. L., et al. 2018, *ATel*, **11608**, 1
- Patterson, J. 1984, *ApJS*, **54**, 443
- Poznanski, D., Prochaska, J. X., & Bloom, J. S. 2012, *MNRAS*, **426**, 1465
- Preibisch, T., Zeidler, P., Ratzka, T., Roccatagliata, V., & Petr-Gotzens, M. G. 2014, *A&A*, **572**, A116
- Rabus, M., & Prieto, J. L. 2018, *ATel*, **11506**, 1
- Ratzloff, J. K., Law, N. M., Fors, O., et al. 2019, *PASP*, **131**, 075001
- Rodrigo, C., Solano, E., & Bayo, A. 2012, *SVO Filter Profile Service Version 1.0*
- Schlafly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, **737**, 103
- Schwab, C., Spronck, J. F. P., Tokovinin, A., et al. 2012, *Proc. SPIE*, **8446**, 84460B
- Shakura, N. I., & Sunyaev, R. A. 1973, *A&A*, **500**, 33
- Shappee, B. J., Prieto, J. L., Grupe, D., et al. 2014, *ApJ*, **788**, 48
- Shara, M. M. 1981, *ApJ*, **243**, 926
- Shara, M. M., Doyle, T. F., Lauer, T. R., et al. 2016, *ApJS*, **227**, 1
- Shore, S. N., Kenyon, S. J., Starrfield, S., & Sonneborn, G. 1996, *ApJ*, **456**, 717
- Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, **131**, 1163
- Smith, N. 2006, *ApJ*, **644**, 1151
- Smith, R. G. 1987, *MNRAS*, **227**, 943
- Stanek, K. Z., Holoiien, T. W.-S., Kochanek, C. S., et al. 2018, *ATel*, **11454**, 1
- Starrfield, S., Iliadis, C., & Hix, W. R. 2016, *PASP*, **128**, 051001
- Steinberg, E., & Metzger, B. D. 2018, *MNRAS*, **479**, 687
- Strope, R. J., Schaefer, B. E., & Henden, A. A. 2010, *AJ*, **140**, 34
- Tamuz, O., Mazeh, T., & Zucker, S. 2005, *MNRAS*, **356**, 1466
- Tokovinin, A., Fischer, D. A., Bonati, M., et al. 2013, *PASP*, **125**, 1336
- Warner, B. 1995, *Cataclysmic Variable Stars*, Cambridge Astrophysical Ser. 28 (Cambridge: Cambridge Univ. Press)
- Wee, J., Chakraborty, N., Wang, J., & Penprase, B. E. 2018, *ApJ*, **863**, 90
- Weiss, W. W., Rucinski, S. M., Moffat, A. F. J., et al. 2014, *PASP*, **126**, 573
- Williams, R. E., Hamuy, M., Phillips, M. M., et al. 1991, *ApJ*, **376**, 721
- Wolf, C., Onken, C. A., Luvaul, L. C., et al. 2018, *PASA*, **35**, e010
- Yaron, O., & Gal-Yam, A. 2012, *PASP*, **124**, 668
- Zacharias, N., Finch, C. T., Girard, T. M., et al. 2012, *yCat*, **1322**, 0