




Ultra-compact dwarfs beyond the centre of the Fornax galaxy cluster: hints of UCD formation in low-density environments

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ABSTRACT

Ultra-compact dwarf galaxies (UCDs) were serendipitously discovered by spectroscopic surveys in the Fornax cluster 20 yr ago. Nowadays, it is commonly accepted that many bright UCDs are the nuclei of galaxies that have been stripped. However, this conclusion might be driven by biased samples of UCDs in high-density environments, on which most searches are based. With the deep optical images of the Fornax Deep Survey, combined with public near-infrared data, we revisit the UCD population of the Fornax cluster and search for UCD candidates, for the first time, systematically out to the virial radius of the galaxy cluster. Our search is complete down to magnitude $m_g = 21$ mag or $M_g \sim -10.5$ mag at the distance of the Fornax cluster. The UCD candidates are identified and separated from foreground stars and background galaxies by their optical and near-infrared colours. This primarily utilizes the $u-i/i-Ks$ diagram and a machine learning technique is employed to incorporate other colour combinations to reduce the number of contaminants. The newly identified candidates (44) in addition to the spectroscopically confirmed UCDs (61), increase the number of known Fornax UCD considerably (105). Almost all of the new UCD candidates are located outside the Fornax cluster core (360 kpc), where all of the known UCDs were found. The distribution of UCDs within the Fornax cluster shows that a population of UCDs may form in low-density environments. This most likely challenges the current models of UCD formation.

Key words: galaxies: clusters: individual: Fornax – galaxies: dwarf – galaxies: evolution – galaxies: star clusters: general.

1 INTRODUCTION

In the late 90s, through the spectroscopic surveys of the Fornax galaxy cluster, Hilker et al. (1999) and Drinkwater et al. (2000a) independently reported the detection of very compact objects at the redshift of the cluster, brighter than globular clusters (GCs) and fainter than compact dwarf galaxies. Since then, many studies have been carried out to investigate the origin of these so-called ultra-compact dwarf galaxies (UCDs; Phillipps et al. 2001). UCDs are larger, brighter, and more massive than typical GCs with typical half-light radii of $10 \leq r_h \leq 100$ pc, luminosities between $-13.0 \leq M_g \leq -10.0$ mag, and masses in a range from 2×10^6 to $< 10^8 M_\odot$ (Mieske et al. 2008; Misgeld & Hilker 2011) with predominantly old stellar populations and a wide range of metallicities (Firth, Evstigneeva & Drinkwater 2009; Janz et al. 2016; Zhang et al. 2018; Forbes et al. 2020; Fahrion et al. 2020a). Moreover, they have on average a dynamical to stellar mass ratio $M_{\text{dyn}}/M_* > 1$ ($M_{\text{dyn}}/M_* = 1.7 \pm 0.2$ for massive UCDs with $M > 10^7 M_\odot$), while for typical GCs, $M_{\text{dyn}}/M_* \sim 1$ (Mieske et al. 2013).

As a result of the studies in the past two decades, two main formation scenarios for UCDs are suggested. In the first scenario, UCDs are the remnant nuclei of tidally disrupted galaxies (Bekki et al. 2003). In this scenario, when a nucleated galaxy with a distinct and high-density nuclear star cluster (NSC) undergoes tidal stripping, it loses most of its stars except the central ones, where the gravitational potential is strong enough to confine stars. Additionally, gravitational interactions induce gas in-fall into the centre of the galaxy and initiate star formation which can change the stellar populations of the nucleus (den Brok et al. 2014; Ordenes-Briceño et al. 2018b; Johnston et al. 2020). In response, the stellar populations and star formation history of the future UCD deviates from the original populations of the nucleus. The detection of tidal features (Voggel et al. 2016a; Schweizer et al. 2018), extended haloes (Evstigneeva et al. 2008; Liu et al. 2020), and asymmetries in the morphology (Wittmann et al. 2016) in addition to an extended star formation history (Norris et al. 2015) and multiple stellar populations (Mieske et al. 2008; Da Rocha et al. 2011) that have been observed in some UCDs, support the stripped nuclei scenario. Moreover, the progenitor's central black hole remains unchanged by the stripping, leading to a higher M_{dyn}/M_* . An SMBH with a mass of 15 per cent on the average can explain the elevated M_{dyn}/M_* (Mieske et al. 2013).

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However, there are cases for which this does not apply or it is not the whole answer since the SMBH is unlikely to be massive enough to explain the observed dynamical to stellar mass ratio (Janz et al. 2015). High-resolution observations (spatial and spectral) of the brightest UCDs found supermassive black holes ($>10^6 M_{\odot}$) in a few cases (Seth et al. 2014; Ahn et al. 2017, 2018; Afanasiev et al. 2018).¹ Alternatives such as variations in the initial mass function (IMF) are proposed to explain the elevated M_{dyn}/M_* (Forbes et al. 2014; Villaume et al. 2017; Haghi et al. 2020; Kroupa 2020).

In the second scenario, UCDs are the outcome of star cluster formation processes, either as the brightest and most massive examples of GCs at the bright end of the globular clusters luminosity function (GCLF; Mieske, Hilker & Infante 2002) or as the result of merging stellar super clusters (Fellhauer & Kroupa 2002). In this case, UCDs are the extension of GCs and share most of their properties. Detailed studies of Milky Way GCs have not yet found any evidence for the presence of a dark matter halo (Mashchenko & Sills 2005; Conroy, Loeb & Spergel 2011; Ibata et al. 2013) or a massive black hole in the centre of GCs (Baumgardt et al. 2003) and their M_{dyn}/M_* is around unity. However, alternatives to the SMBH, like variations in the IMF as mentioned above, or a very concentrated cluster of massive stellar remnants (Mahani et al. 2021) have been proposed to explain the observed elevated M_{dyn}/M_* . While the former can not explain the detected supermassive black hole in some UCDs, the latter predicts an observationally identified central kinematical peak which can be interpreted as an SMBH.

The wide range of properties of the known UCDs and their similarity to both GCs and nuclei of dwarf galaxies suggest that they are formed through multiple and co-existing formation channels (Francis et al. 2012; Norris et al. 2015) and the contribution of the stripped nuclei of dwarf galaxies increases with mass and brightness (Brodie et al. 2011; Pfeffer et al. 2016; Mayes et al. 2020). However, our current knowledge of UCDs is limited by the selection function of the surveys within which UCDs were searched for. The known UCDs were found through spectroscopic surveys around massive galaxies (Norris et al. 2014) or in the cores of galaxy clusters/groups. Therefore, there is not much known about UCDs (or similar objects) in the outskirts of galaxy clusters. In the Fornax cluster, all of the confirmed UCDs are located in the inner 360 kpc (about half of the cluster virial radius). This is because, almost all of the previous spectroscopic surveys of compact sources in the Fornax cluster, covered the core of the cluster (Hilker et al. 1999; Mieske, Hilker & Infante 2004; Bergond et al. 2007; Firth et al. 2007; Firth, Drinkwater & Karick 2008; Gregg et al. 2009; Schubert et al. 2010; Pota et al. 2018). Drinkwater et al. (2000b)² covered a larger fraction of the cluster. However, the data reached $m_g = 20$ mag (80 per cent complete) and were therefore limited to the brightest and most massive UCDs. This magnitude corresponds to $M_g = -11.5$ mag and stellar mass of $10^7 M_{\odot}$ at the distance of Fornax. Drinkwater et al. (2000b) did not find any UCD outside of the cluster’s core (Fig. 1).

Because of the small size of the majority of the UCD/GCs, identification of these objects in images is challenging. To spatially resolve them at low redshifts, ground-based observations in excellent seeing conditions or space-based observations are needed (Richtler et al. 2005; den Brok et al. 2014; Jordán et al. 2015; Voggel, Hilker

& Richtler 2016b). Otherwise, the majority of UCD/GCs appear as point-sources which makes them indistinguishable from foreground stars (Milky Way stars) or distant background galaxies. Traditionally, optical photometry is used to select UCD/GC candidates around galaxies and in galaxy clusters (Taylor et al. 2016; Angora et al. 2019; Prole et al. 2019; Cantiello et al. 2020). Adding information from infrared bands can substantially improve this selection. Muñoz et al. (2014) showed that the compact sources (UCD/GCs) in the Virgo cluster follow a well-defined sequence in the optical/near-infrared colour–colour diagram which can be employed as a tool for identifying them (Liu et al. 2015; Powalka et al. 2017; Cantiello et al. 2018; González-Lópezlira et al. 2019; Liu et al. 2020).

Using the optical data from the Fornax Deep Survey (FDS; Iodice et al. 2016; Venhola et al. 2018) combined with the near-infrared observations of the Vista Hemisphere Survey (VHS) and ESO/VISTA archival data, we aim to identify the possible population of UCDs in the outskirts of the Fornax galaxy cluster within its virial radius (700 kpc Drinkwater, Gregg & Colless 2001). In Cantiello et al. (2018), we derived a list of GC candidates in the cluster using the optical data of the FDS. While having the near-infrared data helps to reduce the contamination, the shallower observations in the near-infrared than the optical limits this study to the UCDs and the brightest GCs.

In the next sections, the data, methods, and results are presented. This paper is structured as follows: Section 2 describes the data and data reduction. Section 3 gives a detailed description of the data analysis, which includes source extraction, photometry, creation of the multiwavelength catalogues of the observed sources and measurement of sizes of the detected sources. Section 4 describes our methodology to identify UCD candidates and the applied machine learning technique. In Section 5, we inspect the performance of our methodology using the ACSFCS catalogue of UCD/GCs (Jordán et al. 2015) and the stars from Gaia DR2 (Gaia Collaboration et al. 2018). We also present the catalogue of UCD candidates in the Fornax cluster, including the spectroscopically confirmed UCDs and the most likely UCD candidates from our analysis. This catalogue supersedes the UCD catalogue presented in Cantiello et al. 2020. We discuss the results in Section 6 and conclude our findings in Section 7.

In this paper, we refer to UCDs as objects brighter than $m_g = 21$ mag or $M_g = -10.5$ mag at the Fornax cluster’s distance. This definition is adopted from Cantiello et al. 2020 (see references therein). This magnitude corresponds to the stellar mass of $4 \times 10^6 M_{\odot}$,³ about the same order of magnitude as ω Centauri (NGC 5139; D’Souza & Rix 2013; Baumgardt & Hilker 2018). Note that we do not adopt any lower limit for the sizes of UCDs. As was discussed here, nowadays, it is believed that the observed UCDs are objects of different origins (Hilker 2006). However, the term ‘UCD’ is still used.⁴ Therefore, our definition based on magnitude (brightness) may include different types of compact objects and does not mean that they necessarily are dwarf galaxies.

Throughout this paper, the term ‘core’ and ‘outskirts’ of the Fornax cluster refers to the area within the virial radius of the cluster, inside and outside the central 1° (360 kpc, \sim half the virial radius) from

³Using the photometric transformations in Jester et al. (2005) and considering the average colour $g-r \sim 0.7$ mag, as is observed for the confirmed UCD/GCs, the magnitude $M_g = -10.5$ mag corresponds to $M_V = -10.9$ mag. Assuming an average $(M_*/L_V) \simeq 2$ (Kruijssen 2008), this magnitude corresponds to the stellar mass $4 \times 10^6 M_{\odot}$.

⁴The term ultra-compact dwarf (UCD) was introduced in the early years of the 21st century since they look like the compact dwarf galaxies, but fainter and more compact.

¹At the time of writing this paper, SMBH is confirmed in five UCDs: four UCDs in the Virgo cluster (M60-UCD1, M59-UCD3, VUCD3, M59cO) and one in the Fornax cluster (Fornax-UCD3).

²This survey led to the discovery of the UCDs and identifying such sources was not planned.

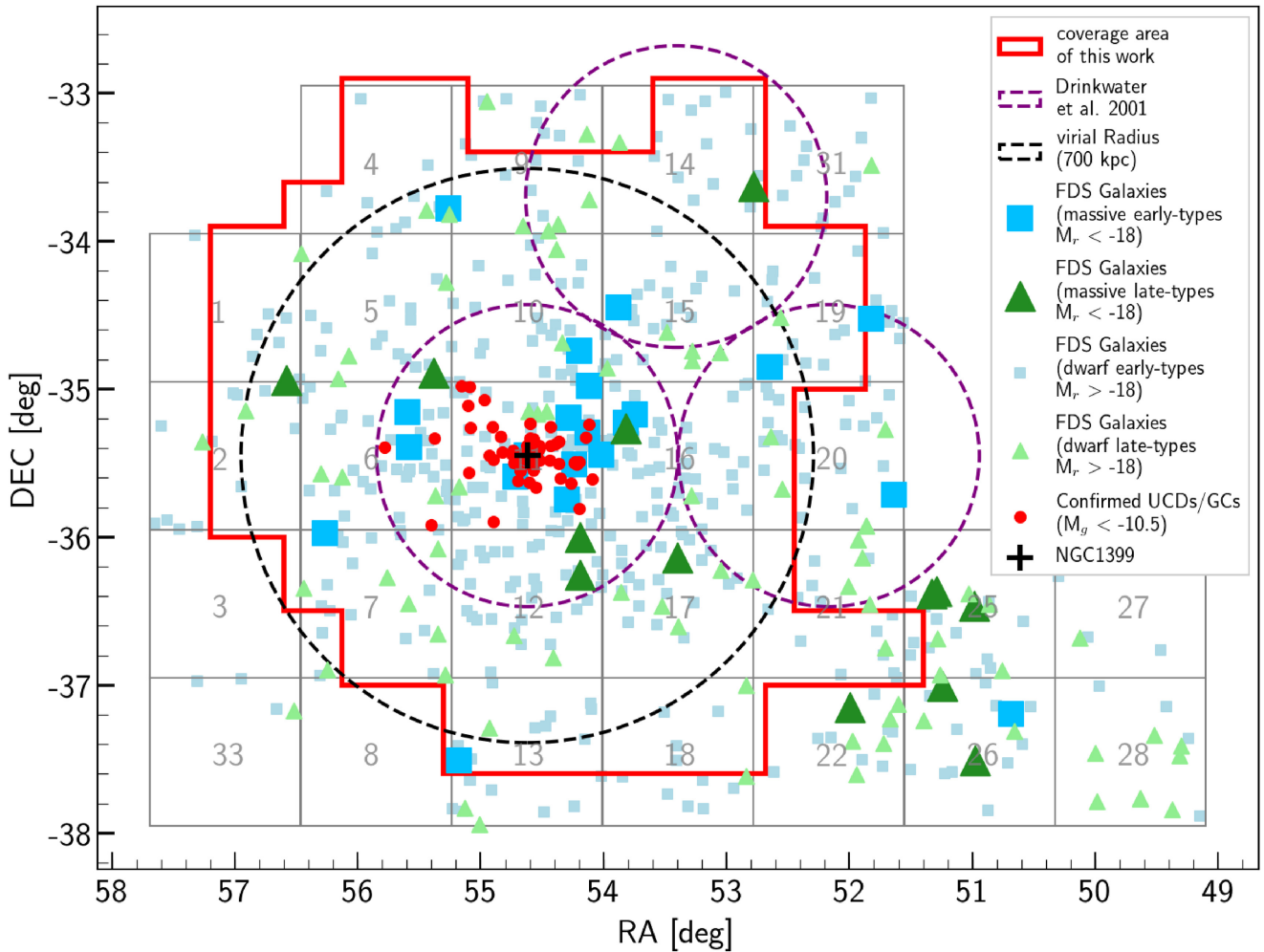


Figure 1. The coverage of the combined optical and near-infrared data in this work. The data set covers the Fornax cluster within its virial radius in six filters (u , g , r , i , J , and K_s) and allows us to search for the compact sources in the cluster outskirts. Most of the previous surveys of compact sources targeted the central 1 deg^2 . As a result, all of the known UCDs (red points) are located in the core of the cluster, within half the virial radius from NGC 1399 in the centre. Drinkwater et al. (2000b) covered a larger fraction of the cluster (purple circles) compared to the other surveys. Grey grid and numbers represent the FDS observed fields.

NGC 1399. Optical and near-infrared magnitudes are expressed in the AB and Vega magnitude systems, respectively. We assume a distance modulus $m - M = 31.50 \text{ mag}$ (a distance of $\sim 20 \text{ Mpc}$) for the Fornax cluster (Jerjen 2003; Blakeslee et al. 2009) and the following cosmological parameters $\Omega_M = 0.3$, $\Omega_\Lambda = 0.7$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

2 OBSERVATIONS

The optical and near-infrared images stem from three different data sets and cover 16 deg^2 out to the virial radius of the Fornax cluster (Fig. 1) in six filters: the optical data of the FDS in u , g , r , i , and near-infrared data from the VHS in J and ESO/VISTA archival data in K_s .

2.1 Optical $ugri$ data: Fornax Deep Survey (FDS)

The FDS (Peletier et al. 2020) is a deep imaging survey using OmegaCAM camera at the ESO VLT survey telescope (VST) and it is used to study a wide range of topics, including intra-cluster light (Iodice et al. 2016, 2017; Spavone et al. 2020), galaxy assembly history (Raj et al. 2020), dwarf galaxies (Venhola et al. 2018 and

Venhola et al. 2019), ultra-diffuse galaxies (Venhola et al. 2017), and GCs (Cantiello et al. 2020). This survey covers 26 deg^2 of the Fornax cluster in four optical bands, u , g , r , and i with an average FWHM of 1.2, 1.1, 1.0, and 1.1 arcsec and a 5σ limiting magnitude (point-source detection with $S/N = 5$) of 24.1, 25.4, 24.9, and 24.0 mag. The depth of this survey reaches the turn-over of the GCs luminosity function at $m_g \sim 24 \text{ mag}$ (Jordán et al. 2007). The optical data is described in more detail in Venhola et al. (2018).

2.2 Near-infrared J data: Vista Hemisphere Survey (VHS)

The VHS (McMahon et al. 2013) is a large survey that covers almost the whole Southern hemisphere ($\sim 20\,000 \text{ deg}^2$) in the near-infrared using the VISTA telescope. This survey is the near-infrared companion of the Dark Energy Survey (DES) and the VST ATLAS survey. VHS provides observations 30 times deeper than 2MASS in J and K_s . We retrieved the reduced frames of the VHS in J (and not K_s) from the ESO science portal.⁵ The final J -band frames have an average FWHM of 1.1 arcsec and a 5σ limiting magnitude of 20.7 mag.

⁵<http://archive.eso.org/scienceportal/>

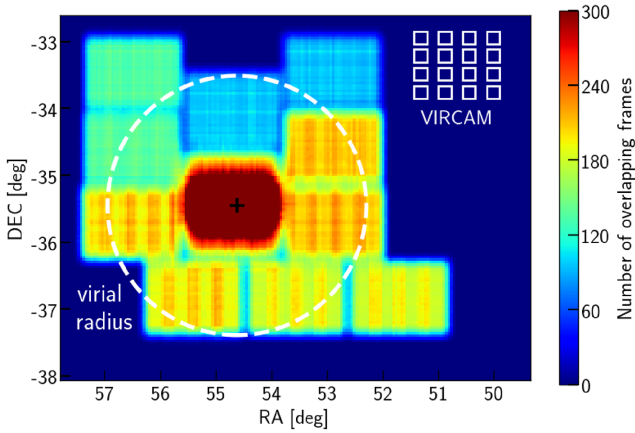


Figure 2. Footprint of the K_s data. The number of overlapping exposures changes across the data, with ~ 300 exposures in the centre (FDS field #11) and ~ 100 – 200 exposures outside of the centre. The black cross in the centre and white dashed circle indicate NGC 1399 and the virial radius of the cluster. The relative sizes and positions of VIRCAM chips (16 chips) projected on the sky are shown in the top right-hand of the figure. The total area covered by pixels (16 chips) in one exposure (paw) is 0.6 deg^2 and six exposures cover a tile of 1.5 deg^2 . The offsets between chips along the x -axis and y -axis are 47.5 and 95 per cent of the size of one chip.

2.3 Near-infrared K_s data: ESO/VISTA archival data

In addition to J band data of the VHS, we used a deeper set of K_s band observations of the Fornax cluster (ESO programme ID 090.B-0.477(A), 092.B-0.622(A), 094.B-0.536(A), 096.B-0.532(A), PI: Muñoz) that were carried out using VIRCAM, the near-infrared camera of the VISTA telescope from 2012 October to 2016 March. VIRCAM consists of 16 chips of 2048×2048 pixel and pixel scale of $0.334 \text{ arcsec pixel}^{-1}$. However, because of the physical distances between the chips, six exposures are needed to cover the telescope’s field-of-view uniformly. This dither pattern covers each piece of sky at least twice (at least 2 pixel of the camera) except for the edges.

The data set consists of 6000 exposures with exposure times of 24 or 30 s and covers the Fornax cluster and partly the Fornax A group in the south-west. The footprint of the data is shown in Fig. 2. When combined, as we show later, these data provide a deeper data set than the VHS K_s data ($\sim 0.7 \text{ mag}$ deeper). The same data is used by Ordenes-Briceño et al. (2018b), who reduced the data independently for the central region of the Fornax cluster (1° around NGC 1399).

The K_s data reduction was performed to remove the instrumental signatures, contaminant sources, and atmospheric effects. Subsequently, the data were calibrated astrometrically and photometrically. The final mosaic was made by stacking all the reduced/calibrated frames. For the data reduction, we used Astro-WISE (OmegaCEN; Begeman et al. 2013; McFarland et al. 2013), an environment developed to analyse wide-field astronomical imaging data in the optical. At first, master-flats were made by median stacking of all the flat-fields taken during each year. Bad-pixel maps were made using the master flats for each chip (16 chips in total). VIRCAM is known to have many bad pixels and bad patches across its chips.⁶ Therefore, making bad pixel maps is necessary to flag these regions through the data reduction. Following this, background frames were made for each science frame. For this purpose, we followed the procedure that has been used in Venhola et al. 2018. After flagging and masking bad

pixels and big/bright objects, science frames were normalized. Then the six science frames that have been taken before and after a science frame within ± 600 s were median stacked to make a background frame. At the end, the background frames were subtracted from the science frames.

The astrometric and photometric calibration of the data was done using the Two Micron All-Sky Survey (2MASS, Skrutskie et al. 2006) as reference catalogue. Astrometric solutions were calculated for individual chips (16 chips) and zero-points were measured on a nightly basis. The photometric calibration on a nightly basis was adopted to increase the number of reference sources and to increase the signal-to-noise ratio (S/N) of the detected sources and achieve a reliable photometry. To do that, all the science frames (reduced and astrometrically calibrated) that were observed during one night were corrected for the atmospheric extinction and then co-added using SWARP (Bertin et al. 2002). 95 per cent of the exposures in K_s have airmass values in a range between 1.0 and 1.3. Considering the K_s extinction coefficient of $0.08 \text{ mag airmass}^{-1}$ (VIRCAM user manual⁷), corrections are $< 0.025 \text{ mag}$.

For each nightly co-added frame, point-sources were extracted and magnitudes were measured using SEXTRACTOR (Bertin & Arnouts 1996). Next, to scale the flux values of all the co-added frames, the SEXTRACTOR MAG_AUTO values were compared to the 2MASS K_s -magnitude, and relative zero-points were calculated. In this step, saturated objects ($m_{K_s} < 13.5$) and faint stars with large uncertainties in their 2MASS photometry ($m_{K_s} > 15.5$) were excluded from the analysis. For the SEXTRACTOR MAG_AUTO, a KRON_FACTOR = 2.5 was used. Using this value of KRON_FACTOR, we expect to capture 94 per cent of the total light (Kron 1980). In that case, 6 per cent of the light is lost, which leads to under-estimating the zero-points. A larger KRON_FACTOR leads to larger apertures, lower S/N and larger uncertainty in photometry. No correction for the missing light and zero-points was applied in this step. This correction was done after making the final mosaics (co-addition), in which, a more accurate approach was used for photometry of the point sources.

The measured relative zero-points were used for the frame by frame (nightly frames) relative calibration to scale flux values in different frames before the final co-addition. Next, the nightly co-added frames and the measured relative zero-points were used to make the final mosaic (using SWARP). We used the bad pixel maps for the final co-addition to create weight maps with zero weight for bad pixels and equal weight for all valid pixels. Next, aperture-corrected photometry was done; and the absolute zero-points were calculated. The aperture-corrected photometry is described in Section 3.

The uncertainties of the photometric reference catalogue (2MASS) for objects fainter than the saturation limit of the K_s observations, make the photometric calibration challenging. Fig. 3 compares the measured K_s magnitudes of the point-sources, extracted from the data, compared to the corresponding 2MASS values. For this comparison, sources were detected and their magnitudes were measured with SEXTRACTOR using BACK.SIZE = 64 pixel and BACK.FILTERSIZE = 3 pixel. Then, bright and not-saturated point-sources within $13.5 > m_{K_s} > 15.5 \text{ mag}$ (to avoid too large uncertainties of the 2MASS photometry and to avoid saturated sources in our K_s data) and $0.6 < \text{FWHM} < 0.9 \text{ arcsec}$ were selected. Note that in this magnitude range, the 2MASS magnitudes have uncertainties larger than 0.1 mag, which is consistent with the scatter in Fig. 3. In this figure, we also explore the colour dependence of the acquired photometric solution. No significant colour dependency is seen in

⁶<http://casu.ast.cam.ac.uk/surveys-projects/vista/technical/known-issues>

⁷<https://www.eso.org/sci/facilities/paranal/instruments/vircam/doc/>

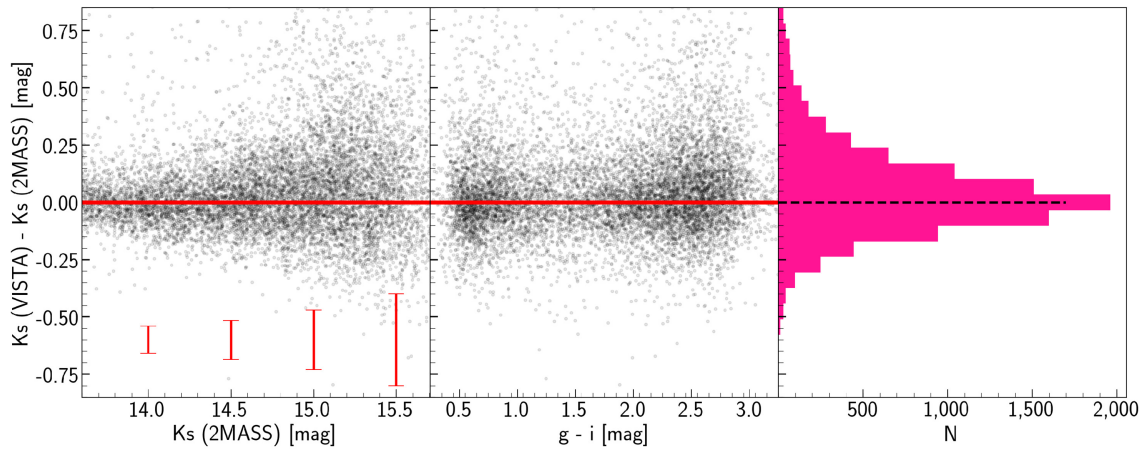


Figure 3. Comparison between the K_s -magnitudes of point-sources in the K_s final (co-added) frame and 2MASS. For this comparison, point-sources within the interval $13.5 > m_{K_s} > 15.5$ mag and $0.6 < \text{FWHM} < 0.9$ arcsec were selected. The scatter in the magnitude residuals is mainly dominated by the uncertainties in the 2MASS photometry. The typical 2MASS uncertainties are indicated by red error bars in the left-hand panel.

Table 1. FWHM values, FWHM standard deviations, and 5σ limiting magnitudes of the K_s data for each FDS field.

FDS field	FWHM (arcsec)	σ_{FWHM} (arcsec)	Limiting magnitude (mag)
–			
2	0.75	0.03	18.5
4	0.73	0.05	18.1
5	0.72	0.04	18.2
6	0.71	0.03	18.5
7	0.72	0.03	18.5
9	0.70	0.05	18.1
10	0.67	0.04	18.3
11	0.68	0.01	19.1
12	0.71	0.03	18.6
13	0.75	0.03	18.4
14	0.70	0.05	18.1
15	0.73	0.05	18.4
16	0.72	0.03	18.7
17	0.71	0.04	18.4
18	0.71	0.04	18.3
19	0.74	0.04	19.0
20	0.75	0.03	18.5
21	0.72	0.04	17.4

the data. Because of the in-homogeneity of the observed regions, the S/N varies across the field. The S/N can drop significantly in some regions, which introduces some low-signal-to-noise regions in the final co-added frame, which cover around 3 per cent of the data.

The final co-added frames⁸ have an average FWHM ~ 0.7 arcsec and a 5σ limiting magnitude of 18.4 mag (~ 1.0 arcsec and 17.7 mag for the K_s data of VHS). The values of FWHM, depth, and limiting magnitude of the final K_s frame are listed in Table 1 for each FDS field.

3 DATA PROCESSING

Because of the small angular sizes of the compact stellar objects (UCD/GCs), identification of this type of object based on size and morphology is challenging. Using the space-based observations of

the ACS/*HST*, in the absence of atmospheric turbulence and using the small pixel size of the ACS camera (0.049 arcsec pixel⁻¹), several thousands of resolved UCDS and GCs have been identified around nearby galaxies and galaxy clusters (Jordán et al. 2004; Sharina, Puzia & Makarov 2005; den Brok et al. 2014; Jordán et al. 2015). However, ground-based observations of compact objects are limited by the seeing. For a typical UCD with half-light radius of ~ 20 pc, assuming a Gaussian profile for the object, the FWHM at the distance of Fornax cluster would be 0.4 arcsec ($\text{FWHM} = 2 \times r_h$). Convolution of this with the typical PSF (Gaussian) FWHM of 1.1 arcsec (for FDS g -band) leads to a final FWHM of 1.17 arcsec. This difference is $1/3$ of the pixel-size of the OmegaCAM (0.21 arcsec pixel⁻¹) and is of the same order as the 1σ uncertainties in g -band FWHM measurements ($\sigma = 0.07$ arcsec or 0.33 pixel). In the case of the K_s data, which provides better image quality than the optical data ($\text{FWHM} = 0.7$ arcsec), the expected FWHM for a typical UCD is 0.1 arcsec (~ 0.3 pixel) larger than the PSF (atmospheric point spread function) FWHM. Note that the near-infrared data has lower S/N than the optical which makes FWHM measurements less certain. As a result, many of the UCDS and all the GCs appear as point-sources, which makes them indistinguishable in shape from foreground stars in the Milky Way or distant galaxies. However, even when they are large enough to be separated from foreground stars, it is still a challenge to distinguish them from slightly resolved background galaxies.

Traditionally, in addition to the sizes of sources, optical colours are used to find UCD/GC candidates around galaxies. However, separating compact sources (UCD/GCs) from the foreground stars solely using the optical colours is very challenging and optically selected candidates can be highly contaminated by foreground stars. Muñoz et al. (2014) showed that the GCs and UCDS of the Virgo galaxy cluster follow a well-defined sequence in the $u-i-i-K_s$ colour–colour diagram (the compact object sequence). This sequence provides a way to identify UCD/GCs and separate them from the foreground stars and background galaxies. However, because most of the time the near-infrared data is shallower than the optical, we can only use this selection for the brightest sources.

3.1 Source extraction

Catalogues of the detected point-sources in each filter (six filters in total: u , g , r , i , J , and K_s) were made using SEXTRACTOR. Because

⁸The final reduced frames can be provided for interested parties upon request.

Table 2. SEXTRACTOR parameters that were used for source extraction and photometry expressed for each filter.

	<i>u</i>	<i>g, r, i</i>	<i>J</i>	<i>Ks</i>
DETECT_MINAREA	5	8	5	5
DETECT_THRESH	1.5	2.0	1.5	1.5
ANALYSIS_THRESH	1.5	2.0	1.5	1.5
DEBLEND_NTHRESH	64	64	64	64
DEBLEND_MINCONT	0.005	0.005	0.005	0.005
PHOT_APERTURES	8,75	8,60	6,30	6,30
BACKPHOTO_TYPE	LOCAL	LOCAL	LOCAL	LOCAL
BACKPHOTO_THICK	20	20	20	20

of the differences in the depth, FWHM and pixel-scales in different filters, different sets of SEXTRACTOR parameters were used for the source extraction and photometry. These parameters are presented in Table 2. The PHOT_APERTURES parameter in the table indicates the sizes (diameter) of the apertures used for aperture-corrected photometry (in pixels) and it is described later in the text.

When running SEXTRACTOR, we used a separate frame for detection. The detection frames were made by subtracting the median filtered science frames using a 16×16 pixel filter. This removes the extended light of the galaxies and increases the efficiency of the source detection around the bright objects except of near their centres (angular distances smaller than 30 arcsec or 3 kpc).

To clean the SEXTRACTOR output catalogues and remove bad (false) detections, such as bad pixels, noise, and cosmic rays, the extracted sources with $\text{FLAGS}^9 > 3$ were removed. In this way, we only include good detections and blended sources. Moreover, detected sources with $\text{FWHM_IMAGE} < 1.0$ pixel were excluded from the catalogues since they are too small to be real and most likely are bad detections or cosmic rays.

Once the single-band catalogues are made and cleaned, they were cross-matched within 1.0 arcsec uncertainty in their positions to create a multiwavelength catalogue. Through this paper, we refer to this catalogue as the master catalogue. The master catalogue includes the multiwavelength photometric information (magnitude, colour, FWHM, ellipticity, etc.) of $\sim 1000\,000$ objects in at least three filters (*g*, *r*, and *i*) and $\sim 120\,000$ objects in all the six filters (*u*, *g*, *r*, *i*, *J*, and *Ks*). For analysis, we only consider the subset with full photometric coverage in six filters and throughout the paper, we refer to it as the main catalogue. Among all the filters, the completeness of sources in the main catalogue is mostly limited by (in order) *Ks*, *u*, and *J*. In total, 88 and 68 per cent of the extracted sources in *g* data brighter than $m_g = 21$ and 22 mag, respectively, are also identified in *u*, *J*, and *Ks* and included in the main catalogue. By assuming that the *g* detections are complete up to $m_g = 22$ mag, this means that the sources in the main catalogue are 88 and 68 per cent complete to $m_g = 21$ and 22 mag.

3.2 Photometry

Good photometry for the faint point sources can be achieved using aperture photometry. However, because of the spatial variation of the FWHM (Fig. 4), aperture photometry with a constant aperture size does not provide *solid results*. In that case, the aperture sizes should be adjusted to FWHM and filter. A solution is to perform photometry

within a constant aperture (using SEXTRACTOR) for each filter and correct the photometry for the applied aperture-size as the following: At first, magnitudes are measured using a smaller aperture (initial aperture) and then, they are corrected for the fraction of missing light. This fraction is estimated by doing aperture photometry for the nearest bright and isolated stars (stars in the same region as the object which have the same FWHM) using two different aperture sizes; the size of the initial aperture and a larger aperture to capture the total light of the bright stars. The difference between these two magnitudes is an estimation for the fraction of the missing light and should be added to the aperture magnitude (using initial-aperture size) of all the objects.

We use the SEXTRACTOR output (aperture magnitude: MAG_APER) to perform aperture corrected photometry for all the sources. At first, we made a star catalogue (real point-sources) based on the second Gaia data release (Gaia DR2; Gaia Collaboration et al. 2018) by selecting the objects with with accurate distance measurements, given by $\frac{\text{parallax}}{\text{parallax_error}} > 5.0$. Afterwards, we used aperture sizes as presented in Table 2. We measured the magnitude difference between the small and large apertures for the stars in the Gaia-based star catalogue and calculated the median of this difference. The calculated value was added to the measured aperture magnitude (with smaller aperture size) of all the sources. To avoid uncertainties caused by FWHM variations, we selected the nearest bright (and not-saturated) stars (closer than 15 arcmin) for each source. The aperture correction values for each pointing are calculated based on 50–100 stars and have an average rms of ~ 0.02 mag.

Once the magnitudes were estimated (using aperture corrected photometry), we re-calibrated the zero-points of the optical data using the overlapping regions between fields as follows. For a given filter, we identified stars in the overlap region between pairs of fields. Next, we measured the average magnitude residuals between the fields using the brightest (non-saturated) overlapped stars. Then, we calculated the absolute value of the zero-point corrections for a given field. The uncertainties associated to the zero-points of the FDS as is estimated by Venhola et al. (2018) are 0.03 mag in *g*, *r*, *i*, and 0.04 mag in *u*. These values are calculated using the stellar locus test (Ivezić et al. 2004). Therefore, since the global zero-points are precise enough, to calculate the (relative) zero-point re-calibration values, we assumed that the average re-calibration value is zero. In other words, the global re-calibration value of the whole data is zero. Table 3 presents the values of zero-point re-calibration corresponding to the FDS fields. All the optical magnitudes from aperture-corrected photometry were corrected accordingly. Next, we address the consistency of the measured magnitudes and colours by direct comparison with the sources in the ACS survey of the Fornax cluster (ACSFCS; Jordán et al. 2015) and KiDS DR4 (Kuijken et al. 2019).

A sample of GCs from the ACSFCS survey (Jordán et al. 2015) is selected based on the high-resolution imaging of the ACS camera of the Hubble space telescope (*HST*), which resolves the GCs larger than 0.01 arcsec (~ 0.2 ACS pixels) or 0.96 pc at the distance of the Fornax cluster. This catalogue provides accurate magnitudes and sizes of sources in two filters g_{475} and z_{850} and GC probabilities p_{gc} , which is the probability that an object is a GC. This catalogue also provides a very accurate (relative) astrometry for the sources. In Jordán et al. (2007), the authors stated that for 35 galaxies out of 43 target galaxies of the ACSFCS, they could find enough astrometric reference sources to derive a reliable astrometry. The galaxies for which they could not carry out the correction were usually the brighter ones.¹⁰ We also

⁹The FLAGS parameter in SEXTRACTOR output is an integer which indicates the confidence level and the quality of the source extraction/photometry for a detection.

¹⁰Jordán (private conversation).

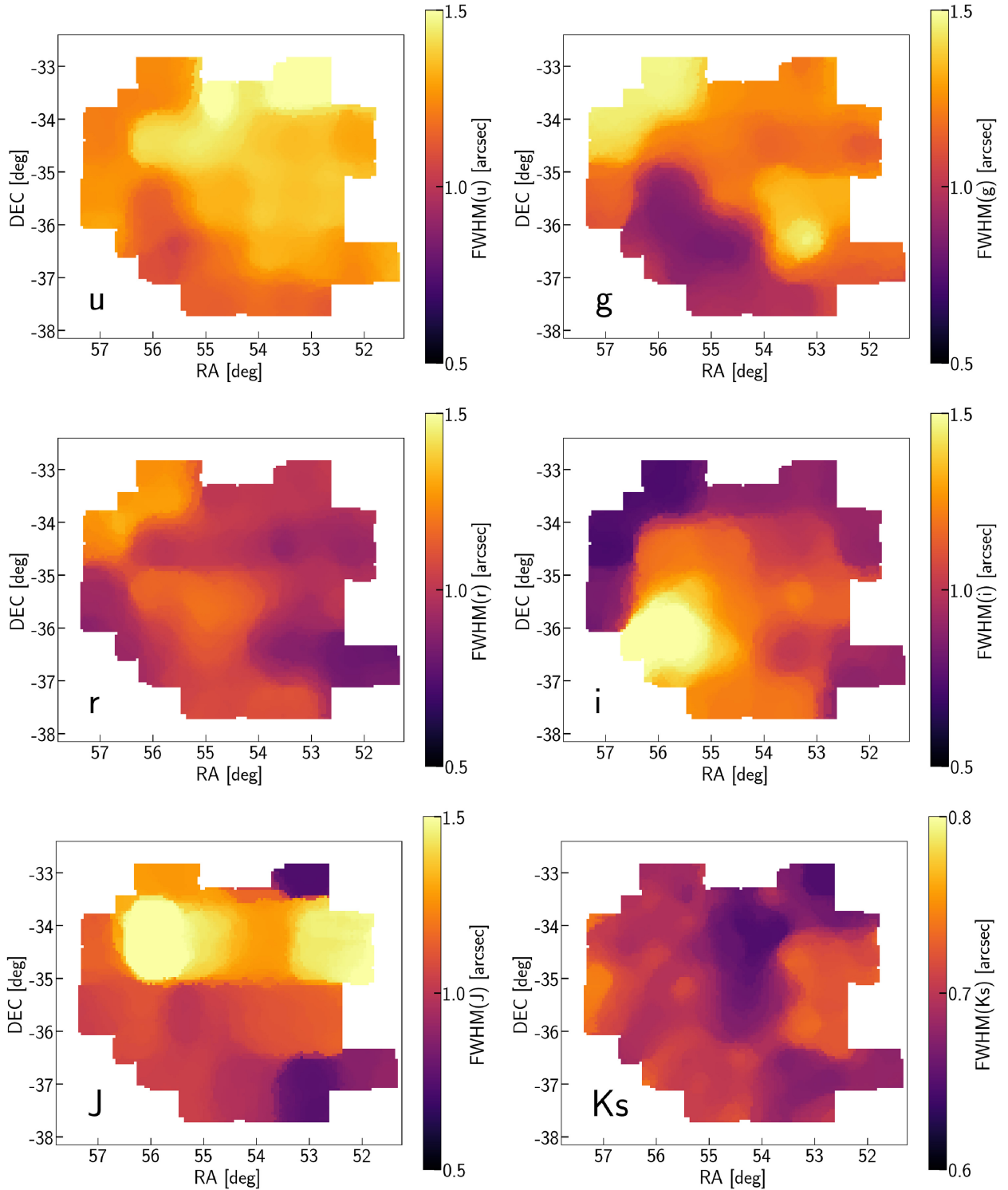


Figure 4. FWHM across the observed field in *u*, *g*, *r*, *i*, *J*, and *Ks*. Among these filters, *Ks* has the best and the most stable FWHM. Note that the colour bar limits for the *Ks* are different than those for the other filters.

found an offset between the coordinates in our data and ACSFCS catalogue mainly around the brightest galaxies. In the most extreme cases, for the objects around NGC 1399 (FCC213) and NGC 1316 (FCC21), while the relative astrometry is good, an >1 arcsec offset in absolute astrometry can be seen. Therefore, before any direct comparison, we corrected the absolute astrometry of the sources. The correction values in RA (Δ RA) and Dec. (Δ DEC) are presented

in Table 4. Δ RA and Δ DEC are the average offset values between our data and Jordán et al. (2015): Δ RA = RA_{FDS} – RA_{ACS} and Δ DEC = DEC_{FDS} – DEC_{ACS}.

We compare the measured *g*-band magnitudes with the *g*₄₇₅ magnitudes of the ACSFCS in Fig. 5. For the comparison, objects with high GC probability $p_{gc} > 0.75$ are selected. The average value for objects brighter than $m_g = 24$ mag is consistent with zero and

Table 3. Zero-point re-calibration values of the FDS fields. The list of fields here represents all the fields with available u , g , r , and i data. Note that the u -band data do not cover fields 22, 25, 26, 27, and 28. Moreover, the u -band data have the largest zero-point re-calibration values. Values ≤ 0.02 are ignored.

FDS field	ΔZP_u (mag)	ΔZP_g (mag)	ΔZP_r (mag)	ΔZP_i (mag)
–	–	–	–	–
1	0.27	–	–	–0.03
2	0.29	–	–	–
4	0.15	–	–	–
5	–0.1	–	0.03	–
6	–0.07	0.07	–	0.06
7	–0.03	–	0.04	–
9	0.07	–	–	–
10	–0.09	–	–	–
11	0.06	0.04	–	0.03
12	–	0.03	0.04	–
13	0.37	–	0.04	–
14	–	–	–	–
15	–	–	–	–
16	–0.17	0.04	–	0.03
17	0.11	–	–	–
18	0.31	–	–	–
19	–0.05	–	–	–
20	–0.05	–	–	–
21	0.15	–	–	–
22	–	–	–	–
25	–	–	–	–
26	–	0.06	–	0.05
27	–	–	–	–
28	–	–	–	–
31	–	–	–	–

rms is $\sigma_{\text{FDS} - \text{ACS}} = 0.15$ mag. For objects brighter than $m_g = 23$ mag with a distance from their host (the nearest bright galaxy) larger than $R_{\text{gal}} = 30$ arcsec (3 kpc) the scatter drops to $\sigma_{\text{FDS} - \text{ACS}} = 0.09$ mag.

Aperture photometry is a reliable method for point-source photometry. However, for extended objects, it misses some of the light, which leads to a larger (fainter) magnitude. Since UCDs can appear slightly extended in the FDS images, we investigated the correlation between the g -band magnitude residuals with the objects' sizes (from ACSFCS). As Fig. 5 shows there is no clear correlation between magnitude difference and size of the objects smaller than $r_h \sim 10$ pc. Fig. 5 also compares the $g-z$ colours of the ACS with the $g-i$ colours from our measurements. Moreover, the stellar loci and colours are consistent with KiDS DR4 (Kuijken et al. 2019) and Covey et al. (2007). Fig. 6 compares the stellar locus in the optical/near-infrared colours space. Stars in our data and KiDS are selected based on the Gaia DR2 parallaxes.

We also compare our magnitudes with the magnitudes of UCDs in Voggel et al. (2016b). These authors, using better seeing conditions (~ 0.5 arcsec) and smaller pixel-size of the instrument (FORIS2/VLT, 0.126 arcsec pixel $^{-1}$) reported the V -band magnitude and half-light radii of 13 UCDs in the Fornax cluster and in the halo of NGC 1399 of which, 12 out of 13 UCDs are identified in our data. These sizes and magnitudes are measured by profile-fitting and take into account the differences in light profiles of UCDs. Fig. 7 compares the V -band magnitudes from our work with those of the full UCD sample from Voggel et al. (2016b), also including UCDs without a size but magnitude measurement. V -band magnitudes m_V (this work) are derived from the photometric transformations in Jester et al. (2005) using $m_V = m_g - 0.58 \times (m_g - m_r) - 0.01$. The consistency of the magnitudes from both works implies that for UCDs with

Table 4. Astrometric offsets of the ACSFCS catalogue of GCs (Jordán et al. 2015).

Host galaxy FCC number	ΔRA (arcsec)	ΔDEC (arcsec)
19	0.53	–0.25
21	0.08	1.30
43	–0.08	–0.23
47	0.17	–0.36
55	0.11	–0.23
83	0.27	–0.37
90	0.11	–0.00
95	0.06	–0.20
100	0.23	–0.22
106	0.25	–0.22
119	–0.13	–0.18
136	–0.19	–0.09
143	0.22	–0.24
147	0.49	–0.84
148	1.11	0.22
153	0.02	–0.03
167	0.09	–0.08
170	–0.05	–0.27
177	–0.13	–0.29
182	–0.08	0.03
184	–0.10	–0.19
190	–0.02	–0.44
193	0.04	–0.32
202	0.10	–0.14
203	–0.14	–0.23
204	0.01	–0.22
213	1.77	–0.10
219	0.39	–0.09
249	0.14	–0.17
255	0.10	–0.22
276	0.05	–0.20
277	0.04	–0.02
288	0.01	–0.17
301	0.04	–0.27
303	0.06	–0.31
310	0.27	–0.13
335	0.25	–0.32

sizes between 20 and 50 pc, the aperture-corrected photometry is reliable.

3.3 Measuring sizes

As discussed above, measuring angular sizes of UCDs at the distance of the Fornax cluster is challenging. Accurate size measurement of compact-sources demands high-resolution space-based observations. However, for the largest UCDs, ground-based observations in an excellent seeing condition (< 0.5 arcsec) are sufficient to measure sizes (half-light radii) as small as ~ 10 pc in the nearby galaxy clusters (Richtler et al. 2005; Voggel et al. 2016b; Liu et al. 2020).

Anders, Gieles & de Grijs (2006) showed that fitting a profile to extra-galactic star clusters, such as a King profile (King 1966) needs high angular resolution and high S/N data. In case of lower S/N data, the authors suggest adopting a simple Gaussian distribution. Following the suggestion in Anders et al. (2006), we assume a Gaussian light profile for the UCD/GCs. In this case, the observed light profile of the objects is the result of a Gaussian distribution (intrinsic distribution) convolved with another Gaussian distribution (atmospheric+instrumental PSF). Therefore, the value of the true

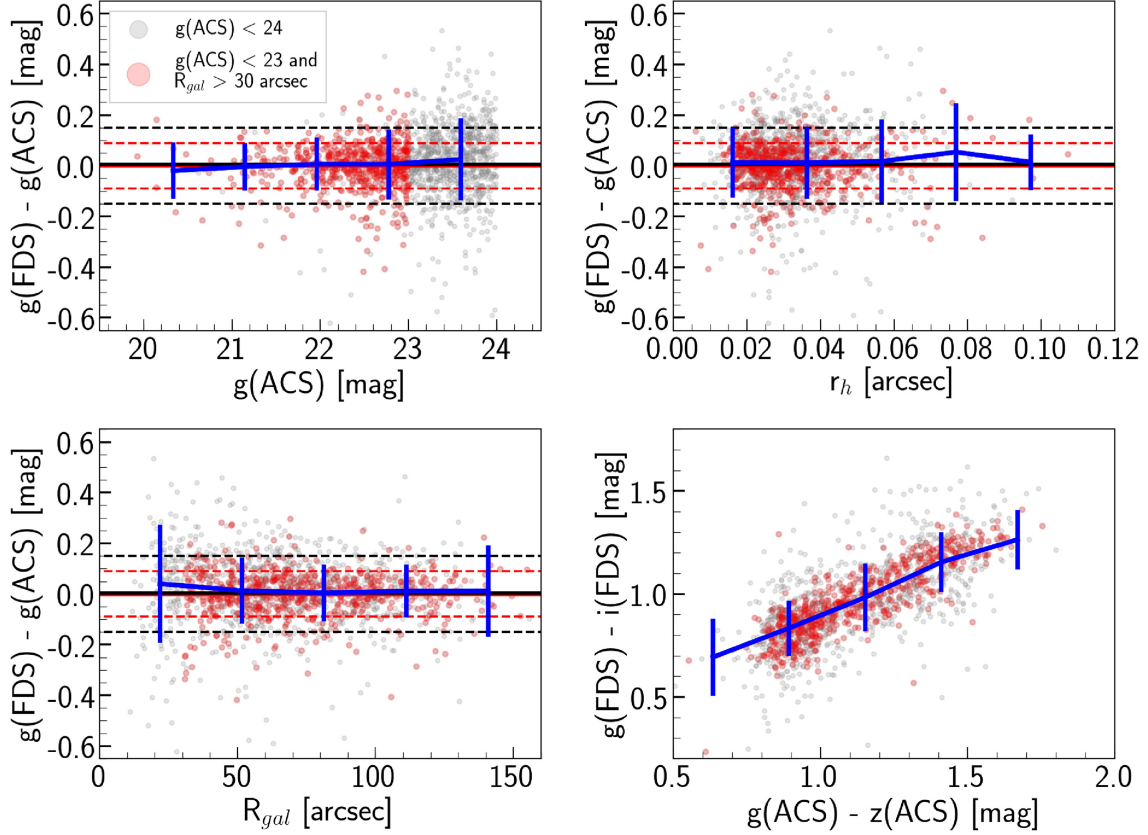


Figure 5. Comparison between our photometry and the ACS survey of the Fornax cluster (ACSFC; Jordán et al. 2015). For the comparison objects with high GC probability $p_{gc} > 0.75$ are selected (from ACSFC). Grey points show sources brighter than $m_g = 24$ mag. Red points represent sources brighter than $m_g = 23$ mag and with $R_{gal} > 30$ arcsec (further than 30 arcsec or 3 kpc from their host galaxy). The solid and dashed lines represent the average and 1σ deviation from average for the corresponding data points of the same colour. For both cases, the average is consistent with zero, and rms are 0.15 and 0.09 mag for grey and red points, respectively. The blue lines and error bars indicate the medians and standard deviations within different bins. The upper right-hand panel shows the correlation between the magnitude residuals and size of the sources (King radius from ACSFC). It is clear that the photometry is consistent for the whole range of sizes.

(intrinsic) FWHM in the absence of atmospheric seeing (or as we call it FWHM*) can be measured using

$$FWHM^* = \sqrt{FWHM^2 - FWHM_{PSF}^2}.$$

Here, FWHM is the measured g -band FWHM of the object (using SEXTRACTOR) and $FWHM_{PSF}$ represents the FWHM of the point spread function (PSF). $FWHM_{PSF}$ is measured as the mean value of the FWHM of point-sources (stars from the Gaia DR2 star catalogue) within 15 arcmin from the object. For objects with $FWHM < FWHM_{PSF}$, $FWHM^*$ is set to zero. The measured $FWHM^*$ value is converted to the half-light radius r_h (also called effective radius R_e) using $FWHM = 2 \times r_h$ (value for a Gaussian distribution).

Using the FDS data with average g -band seeing of 1.1 arcsec, we are able to measure $FWHM^*$ within 0.4 arcsec uncertainties which implies that, for a Gaussian distribution, the uncertainties in measuring the half-light radius of sources is 0.2 arcsec ($FWHM = 2 \times r_h$) or 20 pc at the distance of the Fornax cluster (1 arcsec = 100 pc). In Fig. 7, we make a comparison between our measured sizes and Voggel et al. (2016b). Given the better seeing condition and smaller pixel-size of the instrument of their observations, their reported values are expected to be more precise than ours which provides a benchmark to inspect the accuracy of our measurements. As is seen in Fig. 7 (bottom), our measured values, within the given uncertainties are consistent with those of Voggel et al. (2016b).

For objects smaller than 20 pc, we do not expect to measure reliable sizes. This limit is almost twice as large as the largest UCD/GCs in ACSFC. We use the $FWHM^*$ values in the first step of the UCD identification to exclude the extended sources and we do not define a lower limit on $FWHM^*$ (sizes) of objects. The sizes of the identified UCDs are presented later in the result section.

4 UCD IDENTIFICATION

We use the main catalogue to discover new UCD candidates. We do that in three steps. Here, for clarity, we first shortly summarize these steps.

Step (i). Make the *KNOWN* and *UNKNOWN* catalogues: First, we divide our sources in the main catalogue into two different catalogues, the *KNOWN* and *UNKNOWN* catalogues containing sources with and without available spectroscopic data (radial velocity) in three reference catalogues: Wittmann et al. (2016), Pota et al. (2018), Maddox et al. (2019), and references therein. In the *KNOWN* catalogue, sources are divided into three classes based on their radial velocities. These three classes are: foreground stars, confirmed UCD/GCs and background galaxies.

Step (ii). Pre-selection: Based on the observed magnitude, size, and $u-i-i-Ks$ colours of the confirmed UCD/GCs in the *KNOWN* catalogue, we determine the ranges of their magnitudes, sizes, and

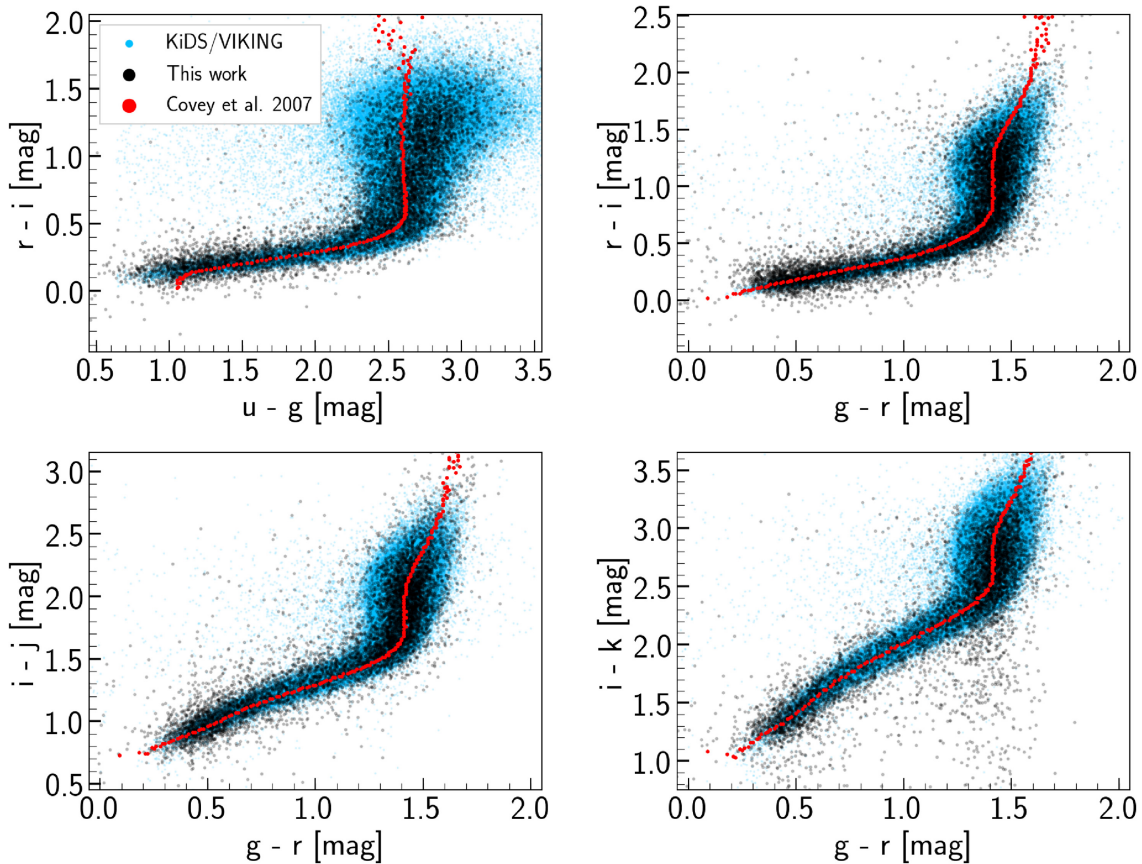


Figure 6. Stellar loci and photometric comparison between our photometry (black dots) and KiDS DR4 (blue dots, Kuijken et al. 2019). On top, the stellar loci from Covey et al. (2007) based on the SDSS/2MASS observations of the bright stars are shown in red.

colours to settle limits on these parameters. In turn we apply these limits to all the objects in the *KNOWN* and *UNKNOWN* catalogues, to exclude likely non-UCDs in the samples. These criteria exclude >95 per cent of the foreground stars and background galaxies (non-UCD/GCs) of the *KNOWN* catalogue while keeping almost all the confirmed UCD/GCs. Therefore, it is also expected that the applied criteria remove the majority (>95 per cent) of the non-UCD/GCs from the *UNKNOWN* catalogue as well.

Step (iii). Selection: we use five independent colours, namely $u-g$, $g-r$, $r-i$, $i-j$, and $J-K_s$ and perform a supervised classification to identify UCD/GCs among the pre-selected *UNKNOWN* sources. It means that we train a machine learning model using the pre-selected objects in the *KNOWN* catalogue (as the training-set) and apply the model to classify the pre-selected sources of the *UNKNOWN* catalogue. By classifying objects, we assign a class (label) to each object in the *UNKNOWN* catalogue. These classes are similar to the three classes of objects in the *KNOWN* catalogue (foreground stars, confirmed UCD/GCs, and background galaxies). Eventually, we select compact sources based on the outcome of the classification as the objects that are classified as UCD/GC. As a result, we have two catalogues: a catalogue of compact sources with $m_g < 21$ mag (UCD candidates in our definition) and one with $21 < m_g < 24.5$ mag (GC candidates).

4.1 Catalogue of *KNOWN* and *UNKNOWN* sources (step i)

The main catalogue has been split into a sample with measured radial velocities in the spectroscopic references (*KNOWN* catalogue) and

the rest (*UNKNOWN* catalogue). In the following, these catalogues are described.

4.1.1 The *KNOWN* catalogue

Sources in the *KNOWN* catalogue (sources with available radial velocities) are categorized into three classes: foreground stars, confirmed UCD/GC, and background galaxies. UCD/GCs are from the spectroscopically (kinematically) confirmed UCD/GCs in the Fornax cluster (Wittmann et al. 2016; Pota et al. 2018; references therein). These UCD/GCs reference catalogues contain 1285 sources in total of which 639 have full photometric coverage and are added to the *KNOWN* catalogue. Table A1 presents the optical and near-infrared magnitudes of these UCD/GCs. The remaining 646 sources lack either the u -band or K_s -band photometry and are not included in the *KNOWN* catalogue.

Foreground stars and background galaxies were selected based on the spectroscopic data of Maddox et al. (2019; references therein). Objects with $V_{\text{rad}} < 300$ km s $^{-1}$ and $V_{\text{rad}} > 3000$ km s $^{-1}$ are selected as foreground stars and background galaxies, respectively. Fornax cluster galaxies have an average radial velocity of 1442 km s $^{-1}$ and a velocity dispersion of 318 km s $^{-1}$ (Maddox et al. 2019). Maddox et al. (2019) presented a compilation of spectroscopic redshifts in the Fornax cluster (Hilker et al. 1999; Drinkwater et al. 2000b; Mieske et al. 2004; Bergond et al. 2007; Firth et al. 2007, 2008; Gregg et al. 2009; Schuberth et al. 2010), also using the extended galaxies of the Fornax cluster Dwarf Galaxy Catalogue (FSDSC; Venhola et al. 2018) to

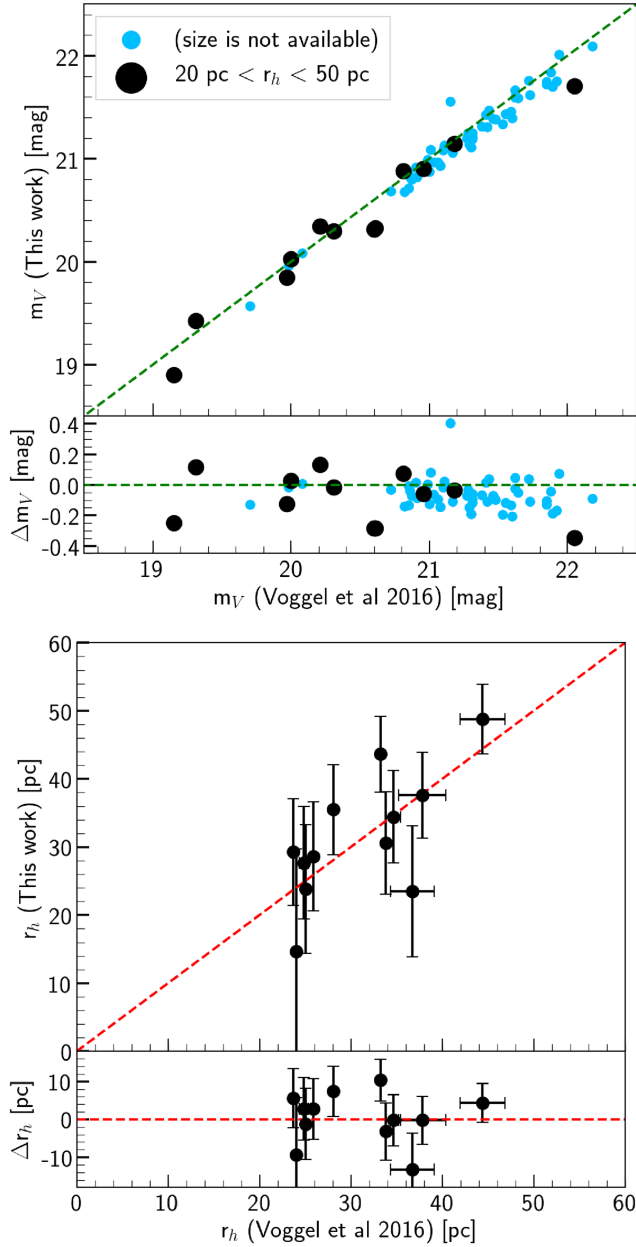


Figure 7. Comparison between the V -band magnitudes (top) and sizes (bottom) of this work and Voggel et al. (2016b) for 12 UCDs (black points). These UCDs are located in the halo of NGC 1399 in the centre of the cluster and have sizes between 20 and 50 pc. Blue points represent the rest of the initial UCD sample of Voggel et al. (2016b) without size estimates since they are too small for model-fitting even in an excellent seeing condition. The uncertainties in magnitudes are smaller than the drawn data points and are not shown in the figure. Δm_V and Δr_h represent the magnitude and size difference between our work and Voggel et al. (2016b).

cross-identify objects. Table A2 and Table A3 gives an overview of the catalogues of foreground stars and background galaxies.

Fig. 8 shows the projected distribution of the *KNOWN* sources on the sky. For convenience, throughout this section of the paper, the spectroscopically confirmed foreground stars, UCD/GCs, and background galaxies (the three classes of objects) in the figures are shown in green, red, and blue, respectively. Almost all the foreground stars in the *KNOWN* catalogue are from the 2dF survey of the Fornax cluster (Drinkwater et al. 2000b); 70 and 25 per cent

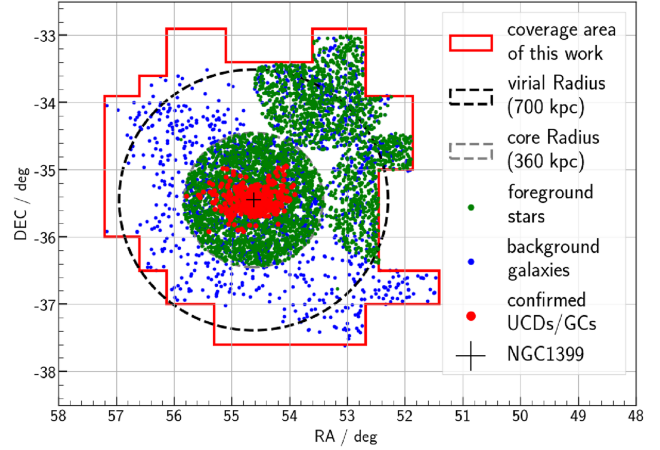


Figure 8. The projected distribution of the sources in the *KNOWN* catalogue. The circular shape of the data-points in the centre is a signature of the 2dF survey of the Fornax cluster (Drinkwater et al. 2000b). The foreground stars (green) and background galaxies (blue) are mainly from two surveys: Drinkwater et al. (2000b; almost all the foreground stars and 70 per cent of the background galaxies) and Maddox et al. (2019; 25 per cent of the background galaxies). Within the covered area, these two surveys are about 80 per cent complete down to $m_g = 20$ mag. The confirmed UCD/GCs (red) are a compilation of many other deeper surveys in the central region of the cluster (Wittmann et al. 2016; Pota et al. 2018; references therein).

of the background galaxies are from the Fornax cluster surveys by Drinkwater et al. (2000b) and Maddox et al. (2019). Both surveys are about 80 per cent complete down to $m_g = 20$ mag. However, the UCD/GCs in the centre of the cluster are a compilation of many other, deeper, surveys. Note that for the *KNOWN* catalogue, we only use sources with available radial velocities in the literature (spectroscopically confirmed). We do not include more stars into the *KNOWN* catalogue from the Gaia DR2 since the spectroscopic sample of foreground stars is rich enough and contains sources over the whole colour range. However, in Section 5, when the UCD candidates are identified, we investigate if any of Gaia DR2 stars are incorrectly identified as UCD.

The *KNOWN* catalogue contains 6670 objects of which 3880 639 and 2151 are foreground stars, UCD/GCs and background galaxies, respectively. This is the first optical/near-infrared data set of the Fornax cluster UCD/GCs and is a valuable source for studying stellar populations of UCD/GCs. This is, however, beyond the scope of this paper and we postpone it to a future publication.

The bottom panel of Fig. 9 shows the magnitude distribution of the sources in the *KNOWN* catalogue. We already discussed that the main catalogue is 88 per cent complete down to magnitude $m_g = 21$ mag. This magnitude corresponds to the adopted magnitude for defining UCDs. We also inspect the completeness of the main catalogue for the red and blue UCD/GCs detected in u and K_s data, respectively. Considering objects with magnitudes $m_g < 21$ mag, $21 < m_g < 21.5$ mag, and $21.5 < m_g < 22$ mag, 100, 96, 83 per cent of the red UCD/GCs ($g - i > 1.0$ mag) are detected in u and 100, 98, 86 per cent of the blue UCD/GCs ($g - i < 1.0$ mag) are detected in K_s .

4.1.2 The UNKNOWN catalogue

The *UNKNOWN* catalogue contains $\sim 110\,000$ objects brighter than $m_g = 24.5$ mag. This magnitude corresponds to the turn over of the GCLF at the distance of the Fornax cluster (Cantiello et al. 2020). The

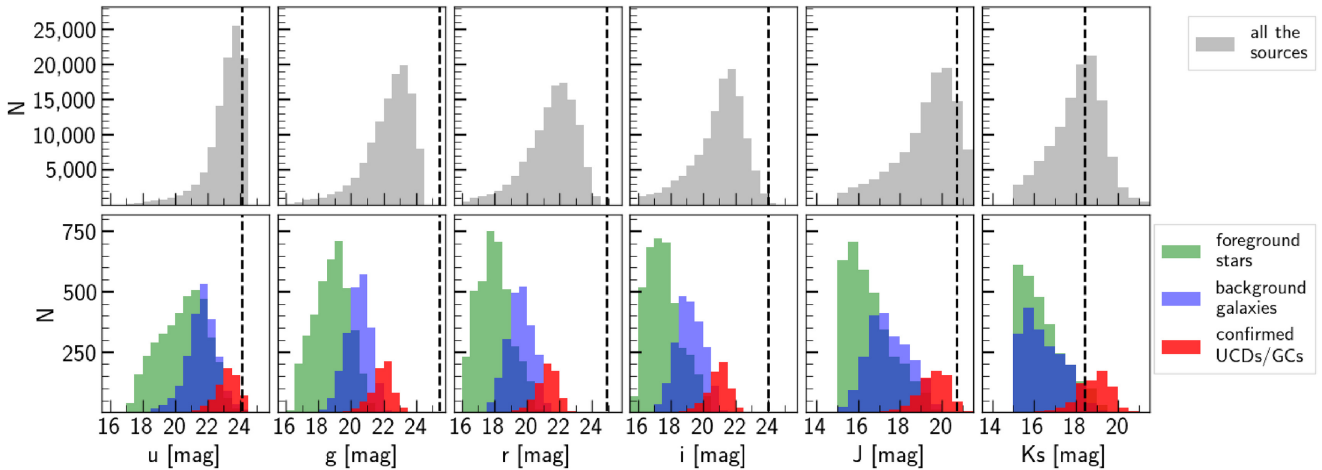


Figure 9. Magnitude distribution of the sources in different filters in the *UNKNOWN* catalogue (top) and the *KNOWN* catalogue (bottom). Foreground stars, confirmed UCD/GCs and background galaxies are shown in green, red, and blue. The 5σ limiting magnitude in each band is indicated by the dashed black line.

UCD candidates will be selected from the sources in this catalogue. Table A4 presents an overview of the *UNKNOWN* catalogue. The top panel of Fig. 9 presents the magnitude distribution of sources in the *UNKNOWN* catalogue in each filter.

4.2 Pre-selection of the UCD/GC candidates (step ii)

In this step, using the observed properties of the confirmed UCDs in the *KNOWN* catalogue, namely g -band magnitude, size, and the $u-i-i-Ks$ colours, we define some criteria to pre-select UCD candidates. The top panel of Fig. 10 shows the $FWHM^*$ -magnitude diagram of the sources in the *KNOWN* and *UNKNOWN* catalogues. All the UCDs (red points brighter than $m_g = 21$ mag) are fainter than $m_g = 19$ and have $FWHM^* < 1.2$ arcsec. By extending these limits in magnitude and size, we adopted a magnitude-size criteria for UCDs as sources with $18 < m_g < 21.5$ mag and $0 \leq FWHM^* < 1.5$ arcsec ($0 \leq r_h < 75$ pc). The bottom panel of Fig. 10 demonstrates these criteria in the $FWHM^*$ versus magnitude diagram. The lower limit on magnitude is 0.5 mag fainter than the magnitude limit of defining UCDs ($m_g = 21$ mag). This decision was made to increase the number of the selected confirmed compact sources, which is important for what comes next in the analysis ($u-i-i-Ks$ colour-colour pre-selection and machine learning). The number of spectroscopically confirmed UCDs in the Fornax cluster brighter than $m_g = 21$ mag is 61. By including fainter compact sources (GCs in our definition) in the magnitude range $21 < m_g < 21.5$ mag, we increase the number of sources by more than a factor of 2–138. We note that the adopted upper limit on $FWHM^*$ does not include Fornax-UCD3 (Drinkwater et al. 2000b), the brightest UCD in the Fornax cluster. This UCD has $FWHM^* = 2.78$ arcsec ($r_h = 139.0$ pc), which is about three times larger than the second brightest UCD in our sample ($r_h = 43.6$ pc). Also, since a few UCDs have $FWHM^* = 0$ arcsec (indistinguishable from 0), we did not define a lower limit on the size. However, the majority of the UCDs (about 90 per cent) have $FWHM^* > 0$ arcsec.

Next, we applied the size-magnitude criteria to the *KNOWN* catalogue and *KNOWN* objects that did not satisfy these criteria were excluded from the rest of the analysis. The total number of size-magnitude selected sources in the *KNOWN* catalogue is 2913 of which 2172, 138, and 603 are foreground stars, UCD/GCs, and background galaxies, respectively. For the *UNKNOWN* catalogue,

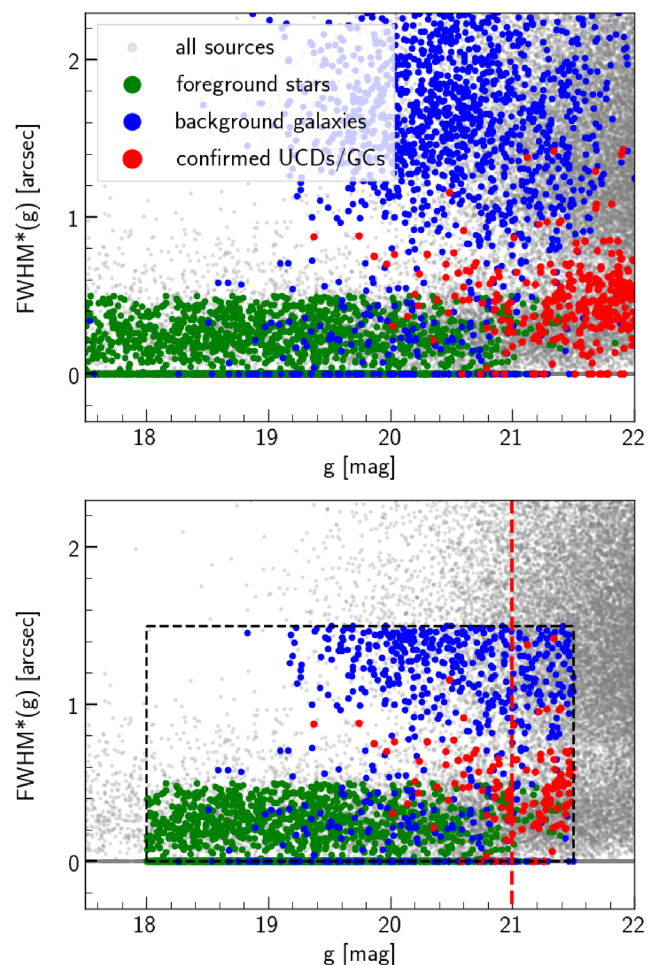


Figure 10. Top panel: $FWHM^*$ -magnitude diagrams (in g band) for sources in the whole data set (grey points), foreground stars (green points), confirmed UCD/GCs (red points), and background galaxies (blue points). The $FWHM^*$ parameter is used as a proxy for the angular sizes of the sources. Bottom panel: The selection boundaries in magnitude and size are shown with black dashed box. The red dashed line indicates our adopted magnitude limit for defining UCDs (Drinkwater et al. 2000b).

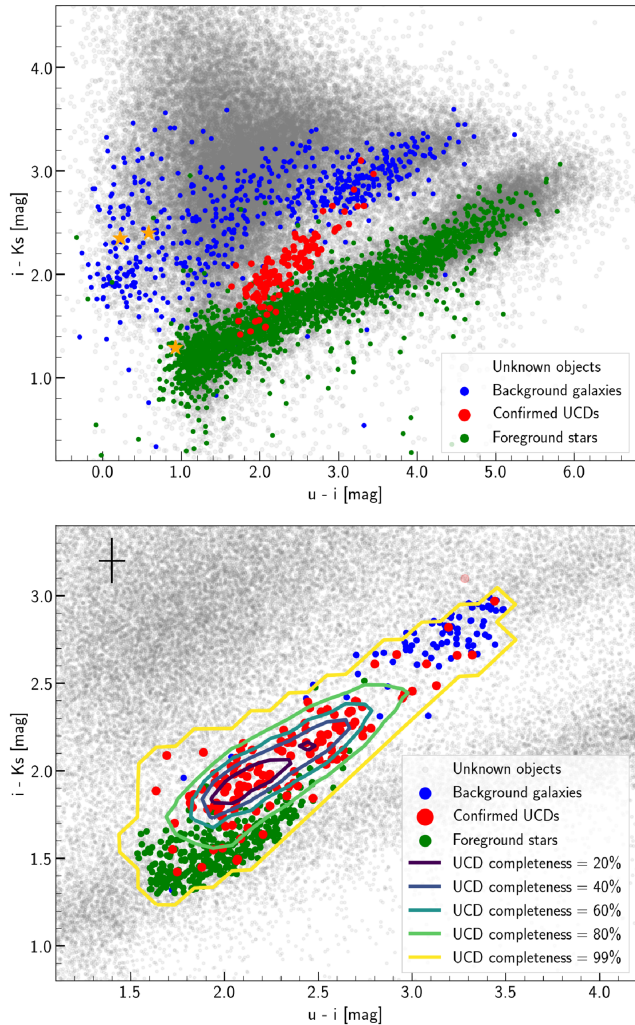


Figure 11. Top panel: The $u-ii-Ks$ colour-colour diagram of the *UNKNOW* sources (grey points), foreground stars (green points), confirmed UCD/GCs (brighter than $m_g = 21.5$, red points), and background galaxies (blue points). The three orange stars indicate the three outliers of the UCD/GC sequence ([BAL2007] gc152.1, [BAL2007] gc290.6, and [BAL2007] gc235.7). These objects were excluded from the analysis. Bottom panel: Pre-selection of UCD/GC candidates using the UCD/GC sequence in $u-ii-Ks$ colour-colour diagram. After the size-magnitude pre-selection, all the *UNKNOW* sources (grey points) within the 99 per cent UCD/GC contour have been selected. Because of the uncertainties in the measurements, the pre-selected sample itself can be largely contaminated with foreground stars (green points) and background galaxies (blue points). The light-red point outside of the UCD/GC sequence indicates a UCD/GC that was excluded after applying the 99 per cent contour. The error bar on top left-hand corner indicates the average uncertainties in colours of a UCD with $m_g = 21$ mag.

we only use the size criteria (and not the magnitude criteria). This is done for two reasons: first, to inspect the quality of our method in rejecting sources brighter than $m_g = 18.0$ as UCDs (which are very likely foreground stars) and secondly, to extend our search to GCs fainter than $m_g = 21.5$ mag. With these criteria, $\sim 60\,000$ objects in the *UNKNOW* catalogue were selected.

We continued the pre-selection using the $u-i$ and $i-Ks$ colours of the sources and the UCD/GC sequence (Muñoz et al. 2014). Fig. 11 (top) shows the $u-ii-Ks$ colour-colour diagram of the size-magnitude selected sources of the *KNOWN* catalogue. In this

figure, UCD/GCs (red points) occupy a region in the colour-colour diagram, which makes them distinguishable from foreground stars (green points) and background galaxies (blue points). The red tail of the sequence is as red as $u-i \sim 3.3$ mag and $i-Ks \sim 3.0$ mag. According to single stellar population models (MILES; Vazdekis et al. 2010, 2016), objects with an old age (10–14 Gyr) and a super-solar metallicity ($[M/H] \sim 0.5$) can be as red as $u-i \sim 3.3$ mag and $i-Ks \sim 2.5$ mag which, within uncertainties are consistent with the observed colours of the red *KNOWN* UCD/GCs in our sample.

The location of the UCD/GC sequence provides a tool to identify UCD/GCs photometrically. To define the UCD/GC sequence in Fig. 11, we used the colours of the confirmed UCD/GCs. The $u-ii-Ks$ colour-colour space was divided into grids with a grid size of 0.1 mag and UCD/GCs’ density was measured within the grids. Then, the isodensity contour enclosing 99 per cent of the UCD/GCs was drawn and UCD/GCs outside this contour were removed from the analysis. As the result, only one object was removed from the initial set of UCD/GCs. We repeated this procedure for the selected UCD/GCs from the first run. No UCD/GC was removed in the second run. Fig. 11 (bottom) shows the corresponding isodensity contours that enclose 20, 40, 60, 80, and 99 per cent of the spectroscopically confirmed UCD/GCs which indicates the UCD/GC completeness of the enclosed colour-colour space. Note that the 99 per cent contour is based on compact sources brighter than $m_g = 21.5$ mag. This contour encloses all the confirmed UCDs (brighter than $m_g = 21$ mag). Once the 99 per cent contour is drawn, objects in the *KNOWN* catalogue (size-magnitude selected) and the *UNKNOW* catalogue (size selected) within this contour were selected and the pre-selection step was completed.

The pre-selection step selects ~ 4079 and 547 objects from *UNKNOW* and *KNOWN* catalogues (331 foreground stars, 137 UCD/GCs, and 79 background galaxies). Table 5 summarizes the numbers of the selected sources in this step. The pre-selected sample, as it is implied from the pre-selected *KNOWN* sources, can be largely contaminated by foreground stars (60 per cent).

The pre-selection based on magnitude, size, and $u-ii-Ks$ colours was done to clean the *KNOWN* and *UNKNOW* catalogues from the obvious non-UCD/GCs (95 per cent of the sources), while it keeps almost all (99 per cent) of the confirmed UCD/GCs. Therefore, the resulting samples are expected to represent a complete sample of all the UCD/GCs while the majority of the foreground stars and background galaxies are removed. The remaining contamination is mainly due to the scatter in the $i-Ks$ colours of the UCD/GCs that originates from the uncertainties in the near-infrared photometry (mainly Ks). Narrowing down the selection area to the densest part of the UCD/GC sequence in the $u-ii-Ks$ diagram (for example 80 per cent completeness contour), while it keeps the majority of the UCD/GCs (80 per cent), it reduces the contamination (85 per cent of foreground stars and 94 per cent of the background galaxies will be removed). However, in this case, a fraction of possible UCDs, in particular, the most metal-rich ones, are missed (20 per cent) and therefore, the resulting sample has a lower completeness. Additionally, for the next step (*step iii*) of the UCD selection (supervised machine learning), the pre-selected *KNOWN* catalogue will be used for the training of the model and it should contain more or less the same number of foreground stars, confirmed UCD/GCs, and background galaxies. Otherwise, the outcome will be biased towards the class with the majority of objects. Our choice with the 99 per cent contour will lead to an equal sample of all three classes. In other cases (for example, using 90 or 95 per cent completeness contours), only a few background galaxies will be included and biases the supervised classification. In the next step, we use all the available colours in the

Table 5. Number of selected objects in *step i* and *step ii*.

Sample	Master catalogue (<i>gri</i>)	Main catalogue (<i>ugriJKs</i>)	Size–magnitude (pre-)selected	<i>u–i/i–Ks</i> (pre-)selected
UNKNOWN	~1000 000	~110 000	~60 000	4079
KNOWN	10 258	6670	2913	547
Foreground stars	5700	3880	2172	331
Confirmed UCD/GCs	1285	639	138	137
Background galaxies	3273	2151	603	79

main catalogue (five independent colours: *u–g*, *g–r*, *r–i*, *i–J*, and *J–Ks*) and try to clean the pre-selected samples.

4.2.1 Outliers from the *u–i/i–Ks* UCD/GC sequence

While inspecting the *u–i/i–Ks* colour–colour diagram of the confirmed UCDs, we found that three spectroscopically confirmed UCDs (brighter than $m_g = 21$) with Simbad ID [BAL2007] gc152.1 ($V_{\text{rad}} = 947 \text{ km s}^{-1}$), [BAL2007] gc290.6 ($V_{\text{rad}} = 1901 \text{ km s}^{-1}$), and [BAL2007] gc235.7 ($V_{\text{rad}} = 1310 \text{ km s}^{-1}$) are located outside of the UCD/GC sequence. Two of these objects ([BAL2007] gc152.1 and [BAL2007] gc290.6) are located in the same region as the quasars and high-redshift objects and one ([BAL2007] gc235.7) is on the stellar sequence. These objects are indicated by orange stars in Fig. 11 (top panel). As discussed in Bergond et al. (2007), the measured radial velocities of these objects are not very certain (flagged ‘B’ in their catalogue). The recent spectroscopic survey of Maddox et al. (2019) updated the radial velocity of [BAL2007] gc152.1 ($\sim 3.8 \times 10^5 \text{ km s}^{-1}$). Therefore, this object is likely a background high-redshift galaxy, as it is expected from its *u–i/i–Ks* colours. Therefore, we conclude that the colours of these three objects do not give enough credibility to them to be a member UCD and they have been excluded from the sample of confirmed UCD/GCs.

4.3 Selection of the UCD/GC candidates (step iii)

In the third step, we use the pre-selected *KNOWN* sources, train a machine learning model using the K-nearest neighbours (KNN) technique with an adjustment to the original technique (described later in Section 4.3.2) and perform a supervised classification to classify *UNKNOWN* sources into three classes: foreground star, UCD/GC, and background galaxy. Our methodology is described, in detail, in the rest of this section. For the classification, we use five features (parameters): five independent colours namely *u–g*, *g–r*, *r–i*, *i–J*, and *J–Ks*. The combination of these colours defines a 5D colour–colour diagram. Fig. 12 shows the different projections of this colour–colour diagram for the pre-selected *KNOWN* sources. Note, we do not use the measured sizes and magnitudes for the classification in this step. Our size measurements are not accurate enough to be used in the classification. Moreover, since the depth of the *KNOWN* catalogues is different for different type of objects (due to different selection criteria and depth in the respective surveys that detected them), including the measured magnitudes in the parameter space biases the classification; for brighter and fainter objects, the classification will be biased towards the foreground stars (which are generally brighter) and UCD/GCs (which are generally fainter).

4.3.1 K-nearest neighbours (KNN)

We aim to find objects that have similar properties as the already known UCDs in the *KNOWN* catalogue. On the other hand, many of the objects are foreground stars or background galaxies which have

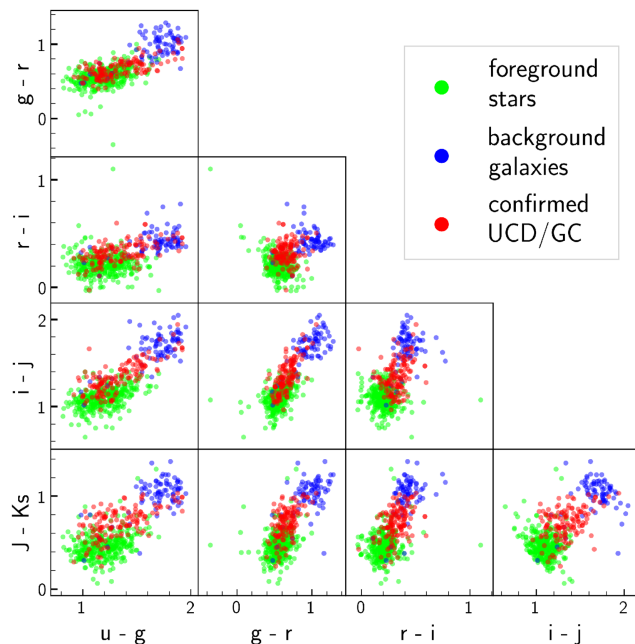


Figure 12. Different projections of the 5D colour–colour diagram of the pre-selected *KNOWN* sources. Foreground stars, UCD/GCs, and background galaxies are shown in green, red, and blue.

properties slightly different from those of UCDs. Using machine learning, we try to identify whether the properties of a given object resemble more those of known foreground stars, background galaxies, and UCDs and label them (classify) accordingly. We use the KNN (Cover & Hart 1967) method, a supervised machine learning technique for classification. Note that our methodology for this step is based on KNN with an adjustment. Here, we briefly explain how KNN works and later, the adopted adjustment to KNN is described.

For each unlabelled object, the KNN algorithm measures its Euclidean distance¹¹ to other data-points and searches for the K-nearest labelled neighbours (called reference-set or training-set) in the parameter space. Based on the classes of the nearest labelled neighbours (foreground stars, UCD/GCs, or background galaxies), KNN classifies (labels) the unlabelled object as the most common class in the neighbourhood. In KNN, the value of K (the number of nearest neighbours) is not arbitrary and it must be optimized based on the efficiency of the classifier using different K values. In any stage of this process, KNN does not apply any normalization to the data.

¹¹In KNN, it is possible to weigh the importance of the neighbours based on their distance in the N-dimensional parameter space, here 5D colour–colour space. During the optimization of the algorithm, we tried weighting and did not find the outcome satisfying. Therefore, for the classification, we did not assign any weight to the neighbours (this can be interpreted as assigning equal weight to all the neighbours).

However, a sort of normalization before running KNN is crucial. Finding the best K value and normalization of the data are described further in Section 4.3.2.

In KNN, the labelled data set is used to make both training-set and validation-set. Normally, the majority of the labelled data is used for the training-set and a smaller fraction is used for the validation-set. Once the classification with the given training-set is done, the validation-set is used to evaluate the accuracy of the classification using the given training-set.

The accuracy of the classifier in classifying a subset (for example, the validation-set) can be expressed by two parameters called precision and recall. The precision describes the reliability of the classifier when it classifies an object as UCD/GC and is defined as the ratio between the number of correctly classified objects as UCD/GCs (true positives) and the total number of objects classified as UCD/GCs (true positives and false positives)

$$\text{Precision} = \frac{\text{True Positive}}{\text{True Positive} + \text{False Positive}}.$$

Recall is the fraction of the confirmed UCD/GCs that are classified correctly (recovered) by the classifier¹² and can be measured as the ratio between objects that are (correctly) identified as UCD/GC (true positives) to the total number of the real UCD/GCs (true positives + false negatives)

$$\text{Recall} = \frac{\text{True Positive}}{\text{True Positive} + \text{False Negatives}}.$$

Precision and recall measure different aspects of a classifier. Another parameter known as F1-score (Rijsbergen 1979), can also be used to express the accuracy of a classifier using one parameter. The F1-score is defined as follows

$$F_1 = \frac{\text{Precision} \times \text{Recall}}{\text{Precision} + \text{Recall}}.$$

Normally, the training and validation are done N times using $\frac{1}{N}$ of the whole data set as validation-set and the remaining $\frac{N-1}{N}$ as training-set. At the end, the average and rms of the precision and recall values for the N runs are the final value and uncertainty of precision and recall of the classifier. This approach is called ‘ N -fold cross-validation’.¹³ The value of K for KNN can be optimized by running the N -fold cross-validation and select the K value which leads to the highest F1-score (the best K value).

4.3.2 KNN+100 and classification

We used the pre-selected *KNOWN* and *UNKNOWN* catalogues as the labelled and unlabelled data sets. Moreover, for our purpose, we adopted a modification to the KNN. For a given training-set, we run KNN for 100 times and each time we randomly select and use 90 per cent of the training-set. Therefore, we classify objects 100 times with a random subset of the training-set. The main motivation behind this modification is to take into account the various outcomes of the classification when a slightly different training-set is used. At the end when classification is done, we count the number of runs that an object is classified as a UCD (UCD_SCORE), foreground

¹²In multiclass classification problems, precision and recall can be measured for the combination of all the classes. However, for simplicity, we measure precision and recall which corresponds to the UCD/GC class.

¹³This method is known as ‘K-fold cross-validation’. However, in this paper, we use ‘N-fold cross-validation’ to avoid confusions with the K value in the KNN method.

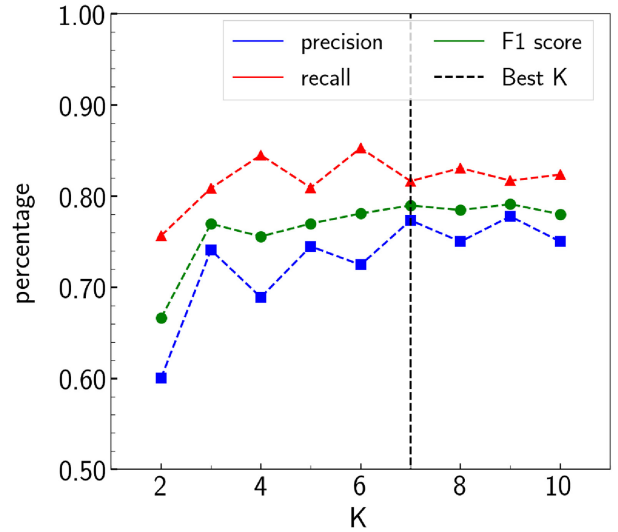


Figure 13. Selecting the best value of K using the estimated F1-score, precision, and recall of different values of K. This result led to the best K value K = 7 for the classification. The values of precision, recall, and F1-score are the average of 10 different runs using 10 different training and validation sets.

star (STAR_SCORE), or background galaxy (GAL_SCORE). These score values are an integer between 0 and 100, and for a given object, the summation of all the scores is 100. The UCD/GC candidates are objects with UCD_SCORE > STAR_SCORE and UCD_SCORE > GAL_SCORE. Clearly, sources with higher UCD_SCORE are more likely to be a UCD. To distinguish the modified version of KNN from the original KNN, we refer to our methodology as KNN+100. Note that in each run, the training-set should contain the same number (same order) of objects of each class. Otherwise, classification is biased towards the class with the majority of the data-points.

Before performing the classification, by visual inspection, we checked that all the confirmed UCD/GCs are in the area outside of the low-signal-to-noise regions of the Ks data. Afterwards, features (i.e. colours) were z -score normalized. z -scores were calculated using $z = (x - \mu)/\sigma$ where μ and σ are the mean and standard deviation of the data-points for a given x . Next, the best K value was derived using N -fold cross-validation with $N = 10$ (10-fold cross-validation) and results are shown in Fig. 13. We chose $K = 7$ as the best K value which leads to precision 77 ± 12 per cent, recall 81 ± 8 per cent and F1-score 79 ± 8 per cent. In practice, there is not much difference between values of K between $K = 5$ and 10. We continued the classification using KNN+100 with $K = 7$ to classify all the sources in the *UNKNOWN* catalogue. The outcome is presented in Section 5.

4.3.3 Performance of the KNN+100 classifier

Fig. 14 presents the result of this classification during 10-fold cross-validation in a confusion matrix. This $N \times N$ matrix shows the three different classes and the number of objects identified in each class. The main-diagonal entries of the confusion matrix show the number (or percentage) of objects classified correctly (true positives) and off-main-diagonal entries show the number (or percentage) of objects classified incorrectly (false positives and false negatives). As is shown in Fig. 14, our classifier can predict (recover) 81 per cent of the UCD/GCs (recall = 0.81). The remaining 19 per cent are classified (incorrectly) as foreground stars (12 per cent) or background galaxies (7 per cent). Moreover, 77 per cent of the predicted UCD/GCs are

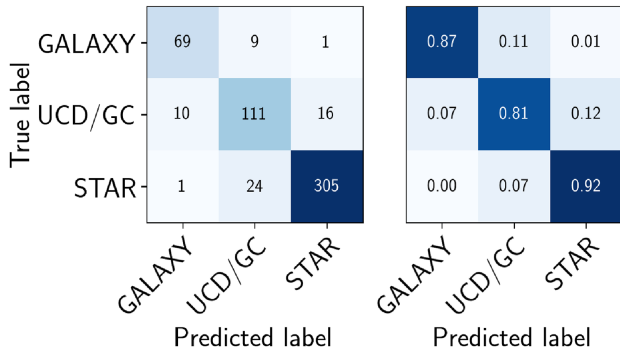


Figure 14. The confusion matrix of the classification for the 10 validation-sets in absolute numbers (left-hand panel) and percentage (right-hand panel). The values in each validation-set can vary by ~ 10 per cent.

true UCD/GC (precision = 0.77), 17 per cent are foreground stars and 6 per cent background galaxies. For the bright sources with $m_g < 20.5$, while precision drops to 42 per cent, recall increases to 90 per cent. For the faint objects with $m_g > 20.5$, precision and recall are 92 and 79 per cent, respectively. Precision and recall for both bright and faint sources improve when objects inside the 80 per cent UCD completeness contour (in $u-i-i-Ks$ diagram) are considered. In this case, for bright sources, precision and recall are 55 and 100 per cent, respectively. For faint sources, these rates are 94 and 89 per cent. Therefore, to have a cleaner sample of UCDs, one might restrict the candidate selection to a lower UCD completeness contour, for example, the 80 per cent completeness contour. The lower rate of precision of the bright sources compared to the faint sources is mainly because of the smaller number of UCDs at brighter magnitudes, while the number of total sources (and therefore the expected false positives) are the same.

Fig. 15 shows the $u-i-i-Ks$ colour–colour diagrams (top) and UCD_SCORE values (bottom) of the true-positives and false-positives (TP and FP, on the left-hand) and true negatives and false-negatives (TN and FN, on the right-hand) in all the 10 validation-sets. The performance of our classifier is bounded to the data points in the training-set and for this reason, except of a few objects, it misses the majority of the confirmed UCD/GCs outside of the 80 per cent UCD completeness contour (false negatives in the figure on right-hand, 10 per cent of the total UCD/GCs). The 80 per cent contour, while it encloses 80 per cent (completeness) of the UCD/GCs, selects 13.5 per cent of the non-UCD/GCs (15 and 6 per cent for the foreground stars and background galaxies, respectively) and has the highest rate of selecting UCDs over non-UCDs. The middle panels of Fig. 15 present the UCD_SCORE of the sources inside the 80 per cent UCD completeness contour. Except of two objects, all the confirmed UCD/GCs have UCD_SCORE > 80 . Only 2 out of 119 UCDs within the 80 per cent contour have UCD_SCORE < 80 which corresponds to 1.6 per cent of the sample. Also, 14 foreground stars and 5 background galaxies within the 80 per cent contour (19 in total) have UCD_SCORE > 80 which corresponds to 14 per cent of the total sources identified as UCDs.

Normally, in addition to constraints on size and magnitude, the $u-i-i-Ks$ colour–colour diagram is used to select UCD/GCs (as we did in *step ii* of the UCD selection). However, as we showed, this criterion alone leads to a contaminated sample (the pre-selected sample) by foreground stars and background galaxies. This is the case especially when the near-infrared data is not as deep as the optical. This is the motivation behind *step iii* in which we cleaned the pre-selected sample from the possible contaminants using all

the available colours (five independent colours) in our data set. This leads to the removal of 70 per cent of the contaminants (64 per cent for the sources inside the 80 per cent contour). However, for objects outside the 80 per cent contour, our methodology is not reliable. This is because our UCD/GC sample does not provide enough data points outside the 80 per cent contour. We also discussed that this contour has the highest rate of selecting UCDs over non-UCDs. To summarize, when objects inside the 80 per cent contour are considered, our methodology is more reliable. Moreover, we showed that among the UCDs inside the 80 per cent contour, almost all the UCDs have UCD_SCORE > 80 . In the next section and after presenting the result of classification of the *UNKNOWN* sources, we select the most likely UCD candidates as objects inside the 80 per cent contour and with UCD_SCORE > 80 . We call these candidates as the ‘BEST’ UCD candidates.

5 RESULTS

We initially identified sources with full photometric coverage in six filters (main catalogue) and divided it into a *KNOWN* and an *UNKNOWN* catalogue, respectively for the objects with and without available radial velocity data in our reference literature (Wittmann et al. 2016; Pota et al. 2018; Maddox et al. 2019). Then, a sample of UCD/GC candidates was pre-selected using size–magnitude criteria and the $u-i-i-Ks$ colour–colour diagram. Subsequently, and by applying a supervised machine learning technique trained on the pre-selected sources in the *KNOWN* catalogue, we classified all the pre-selected sources in the *UNKNOWN* catalogue into three classes: foreground stars, UCD/GCs, and background galaxies. In this section, the results of the classification and the final catalogue of the UCDs are presented.

5.1 UCD/GC candidates

Our method identifies 1155 *UNKNOWN* UCD/GCs brighter than $m_g = 24.5$. By defining the UCDs as objects brighter than $m_g = 21$ mag ($M_g = -10.5$ mag at the distance of the Fornax cluster), we identified 222 UCD candidates and 933 GC candidates in total. The UCD/GC candidates are objects with UCD_SCORE $> STAR_SCORE$ and UCD_SCORE $> GAL_SCORE$. Hence, all the UCD and GC candidates have UCD_SCORE ≥ 50 . A preview of the published catalogue of UCD and GC candidates in this paper is presented in Table A5 (in the appendix). The observed area of this work covers the Fornax cluster up to its viral radius and a bit beyond. Fig. 16 shows the UCD/GC candidates and the recovered UCD/GCs from the *KNOWN* catalogue (true positives) around some of the brightest Fornax cluster galaxies.

Among the UCD/GC candidates, four UCD candidates and five GC candidates (nine in total) have radial velocities in other references in the literature (different than our references in this paper) and are not included in the *KNOWN* catalogue. Among these UCD candidates, one is a foreground star, one is a UCD (this means that one UCD is missed from the *KNOWN* catalogue), one is a nucleated dwarf galaxy, and one is a background galaxy. All these sources have UCD_SCORE = 100. For the five GC candidates, one is a foreground star (UCD_SCORE = 83), two are GCs (UCD_SCORE = 100 for both), one is a nucleated dwarf galaxy (UCD_SCORE = 60), and one is a background galaxy (UCD_SCORE = 100). This implies that 55 ± 25 per cent of the identified UCD/GCs are a true UCD/GC. Note that the nuclei of dwarf galaxies can also be considered as a true UCD since they share similar properties.

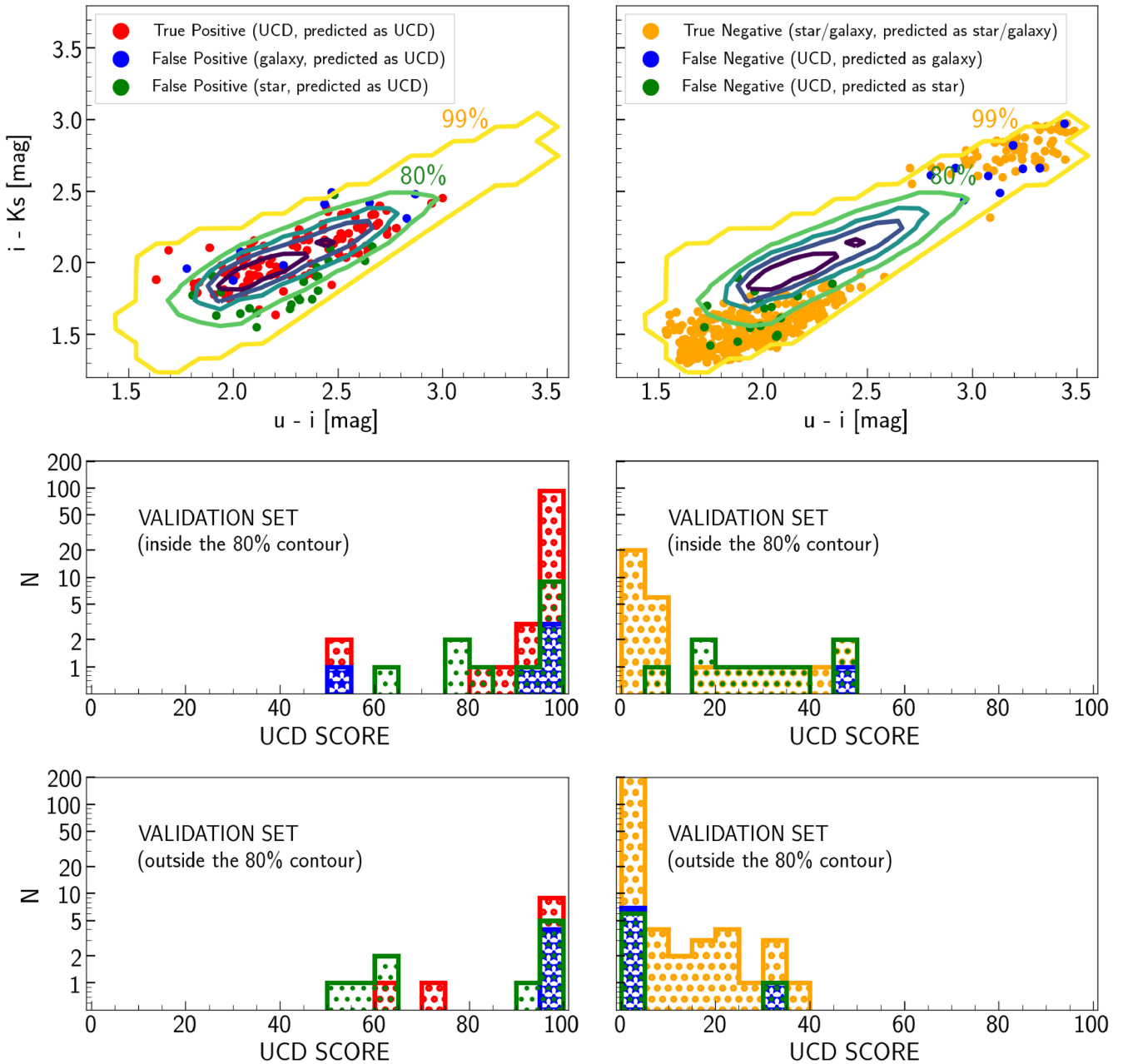


Figure 15. The colour–colour diagram and UCD-SCORE values of the true-positives (TP), true-negatives (TN), false-positives (FP), and false-negatives (FN) in the 10 validation-sets used for 10-fold cross-validation.

Out of about 3000 size–magnitude selected foreground stars and background galaxies in the *KNOWN* catalogue, our classifier ended up classifying about 30 as UCD/GCs (false positives). This implies an average false-positive rate of 1 per cent. In Fig. 17, we show the projected distribution of the selected UCDs and GCs in comparison with the expected false positives derived from the total number of sources in the *UNKNOWN* catalogue. We estimate ~ 170 (~ 77 per cent) false positives and ~ 50 (~ 23 per cent) true UCDs among the UCD candidates. For the GC candidates, we estimate ~ 580 (63 per cent) false positives and ~ 350 (37 per cent) true GCs. Since this is unsatisfactory, we now choose a subsample, with higher probabilities of being UCDs or, as we call it, the BEST UCD candidates.

5.2 BEST UCD candidates

Once the UCD candidates are identified, we select the most likely UCD candidates among them as follows. At first, among the candidates brighter than $m_g = 21$ mag (222 sources), *UNKNOWN* objects inside the 80 per cent contour in the $u-i-i-K_s$ diagram and with high a UCD probability, estimated by the machine learning technique ($UCD_SCORE > 80$) were selected. Next, for this sample, objects with $FWHM^* = 0$ are excluded. The criterion was applied because the majority of the confirmed UCDs (about 90 per cent) have a non-zero $FWHM^*$. Also, by visual inspection, we cleaned the data from blends and objects in the low-signal-to-noise regions of the K_s images. Applying the criteria on the completeness, UCD-SCORE, $FWHM^*$, and finally the visual inspection, we excluded 88, 27, 43,

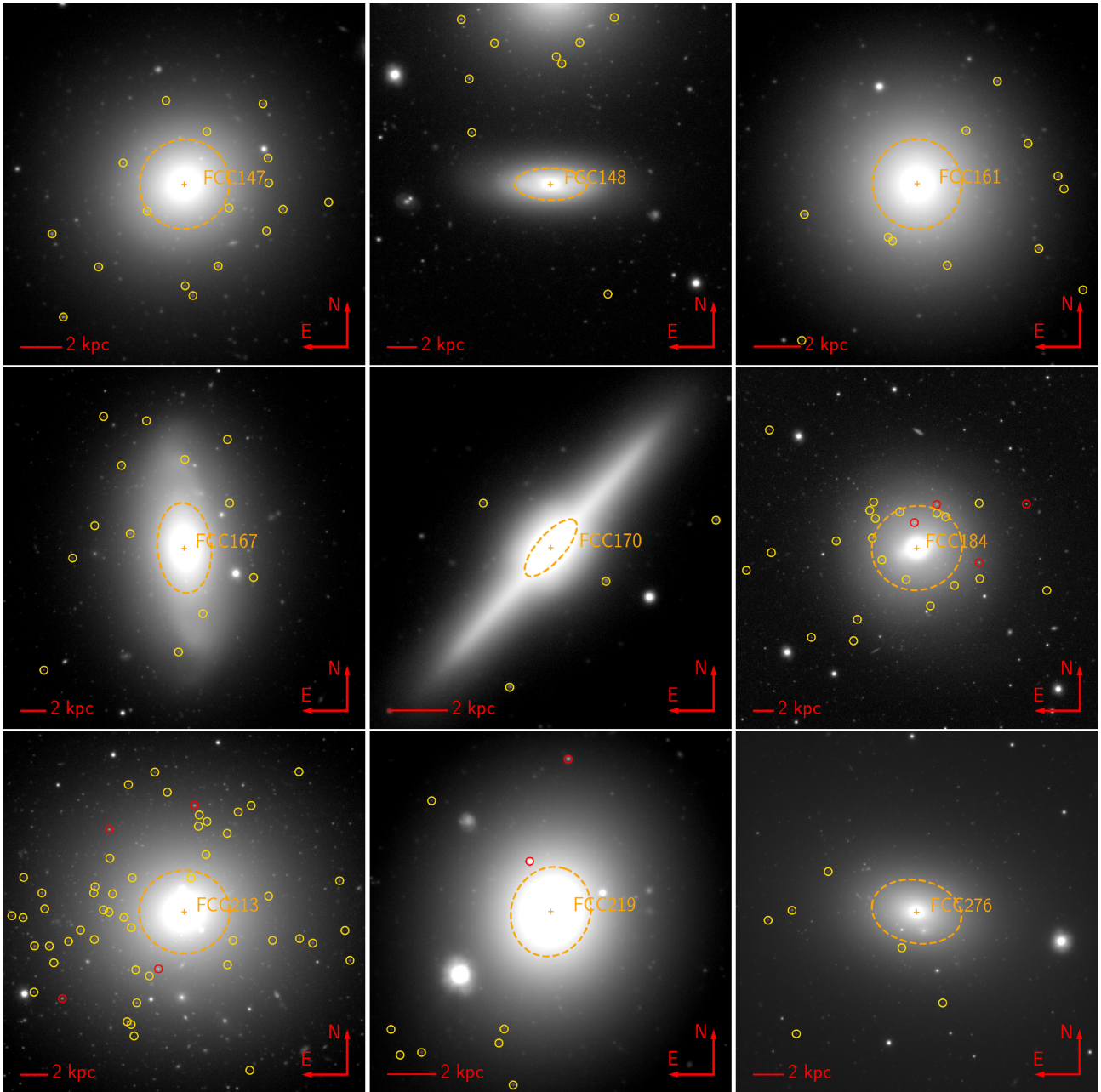


Figure 16. UCD/GCs around the brightest galaxies in the Fornax cluster. These objects represent all the sources that are identified from the sources in the *UNKNOWN* catalogue and the *KNOWN* catalogue (recovered UCD/GCs from the *KNOWN* catalogue or in other words, true positives in all the validation-sets). UCDs and GCs are indicated with red and yellow circles, respectively. The dashed circles indicates one effective radius from the host galaxy (based on the values in the FCC catalogue; Ferguson 1989). Fornax cluster galaxies with two or more identified UCD/GC candidates are shown here. Note that the inner 3 kpc was excluded from the detection of UCD/GC candidates.

and 15 sources (173 in total) from the 222 sources, respectively. These selection criteria leave us with 49 BEST UCD candidates. Applying the same criteria on the validation-sets leads to a UCD sample with completeness of 70 per cent (43 UCDs out of 61 UCDs). It also improves the rate of true positives by a factor of 2 from 23 to 52 per cent.

Among these 49 selected candidates, three are nuclei of dwarf galaxies¹⁴ in the cluster (FCC numbers FCC274, FCC188, and FCC

260), one is a UCD around FCC219 (Fornax 3D survey; Iodice et al. 2019; Fahrion et al. 2020b), one is a star (Schuberth et al. 2010). In the following, all these sources were excluded from the BEST UCD sample and the final number of the BEST UCDs is 44. This number is within the expected range of the number of true positives (true UCDs) among all the UCD candidates (~ 50). The $u-i/i-Ks$ colour-colour diagram of the most likely UCD candidates or ‘BEST’ UCD candidates is shown in Fig. 18.

Fig. 19 shows the thumbnails of the BEST UCDs; and the final catalogue is provided in Table A6. In Figs 20 and 21, we compare the observed properties of the BEST UCD candidates with the confirmed

¹⁴Because of our methodology in source extraction, in addition to the size selection, we miss further nuclei of dwarf galaxies in the final catalogue.

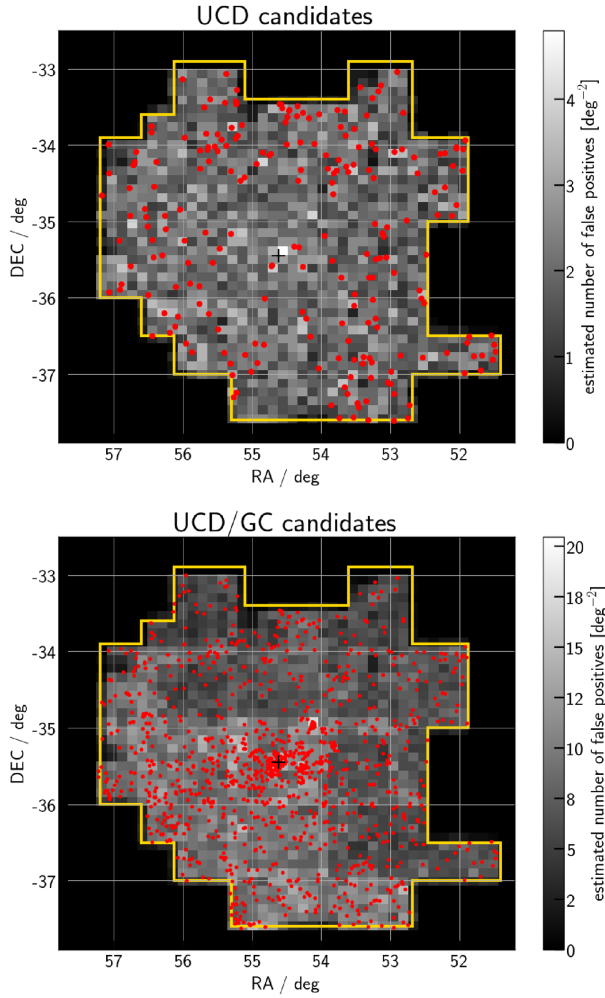


Figure 17. Projected distribution of the candidates and the expected number of the false positives across the observed area for UCD candidates (top panel) and all the UCD/GC candidates (bottom panel).

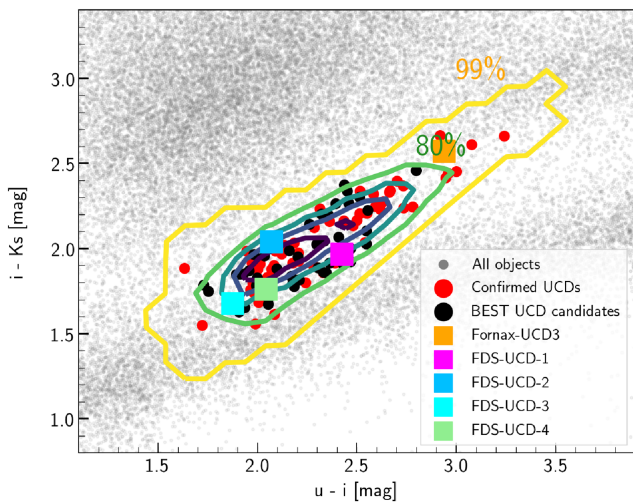


Figure 18. $u-i/i-Ks$ colour-colour diagram of the confirmed UCDs (red) and the BEST UCD candidates (black). The four brightest BEST candidates brighter than $m_g = 19$ mag (FDS-UCD-1, FDS-UCD-2, FDS-UCD-3, FDS-UCD-4) are indicated in the diagram.

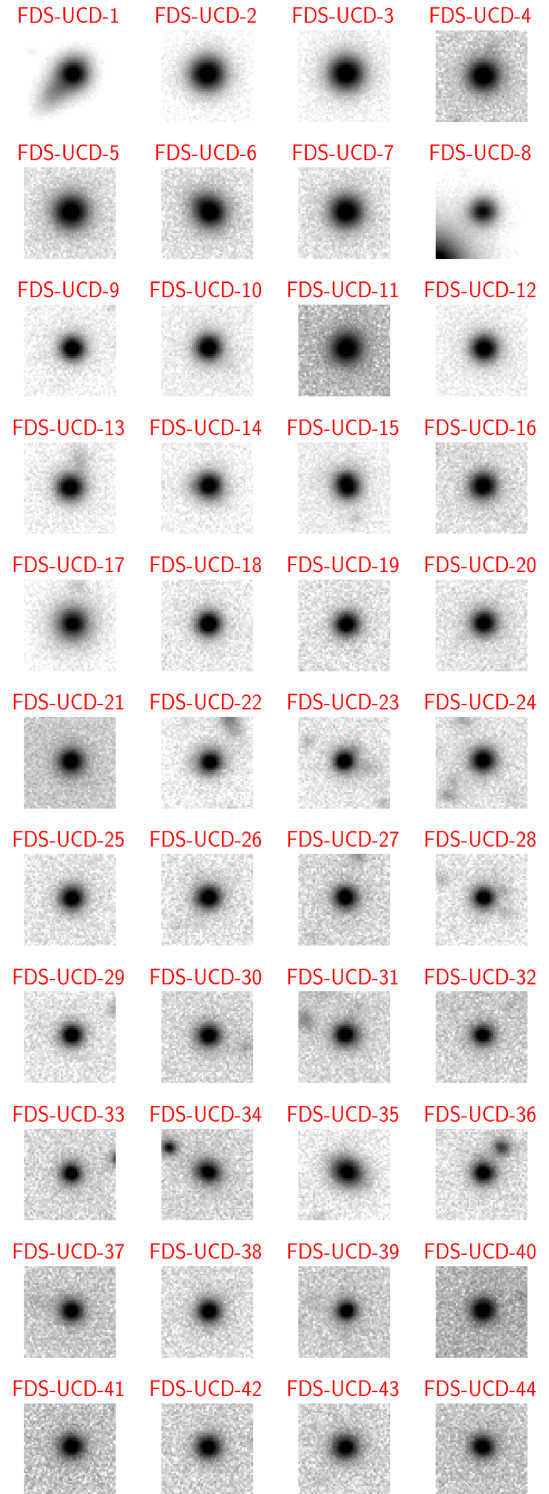


Figure 19. The BEST UCD (candidates) identified in this work. Thumbnails have dimensions of 10 x 10 arcsec or 1 x 1 kpc at the distance of the Fornax cluster. This list is sorted based on m_g . The first four BEST candidates (first row) are the brightest candidates, brighter than $m_g = 19$ mag. An elongated feature can be seen around FDS-UCD-1 which looks to be a blended with a background galaxy.

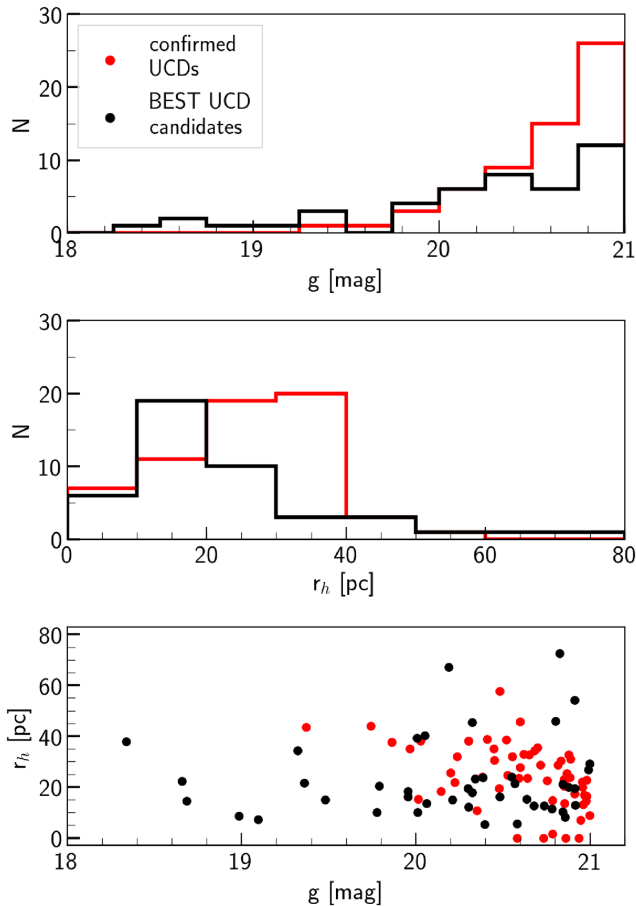


Figure 20. Magnitude distribution, size distribution, and size–magnitude relation of the BEST UCD candidates (in black) and the confirmed UCDs (in red). This figure does not show Fornax-UCD3 in the Fornax cluster. This object is located outside of the displayed size limits ($r_h = 139$ pc based on our measurements).

UCDs. Since almost all of the BEST candidates are outside of the core while all the confirmed UCDs are inside the core of the cluster, the properties of these two samples represent the properties of UCDs in their respective environments. As is seen in Fig. 20, we have identified a few BEST UCD candidates brighter than the brightest confirmed UCD in our pre-selected sample. Among the brightest candidates, four are brighter than $m_g = 19$ mag (FDS-UCD-1, FDS-UCD-2, FDS-UCD-3, FDS-UCD-4). These candidates in the $u-i/i-Ks$ colour–colour diagram are indicated in Fig. 18. Note that, as was discussed earlier, for bright magnitudes, the rate of false positives is higher. The thumbnails of these objects can be seen in the first row of Fig. 19. FDS-UCD-1 is brighter than Fornax-UCD3 (Drinkwater et al. 2000b). In Fig. 19, FDS-UCD-1 seems to be blended with a background galaxy. This blend may have led to a larger measured size, brighter magnitude, and redder colours for this object.

5.3 ACSFCS UCD/GCs

As an external check, we compared our UCD/GC candidates with the GC catalogue of the *HST*/ACS survey of the Fornax cluster (ACSFCS; Jordán et al. 2015). For the comparison, we combined the selected UCD/GCs in the *UNKNOWN* catalogue with the true positives. These true positives are objects in the 10 validation-sets that are recovered as UCD/GCs when they were not included in

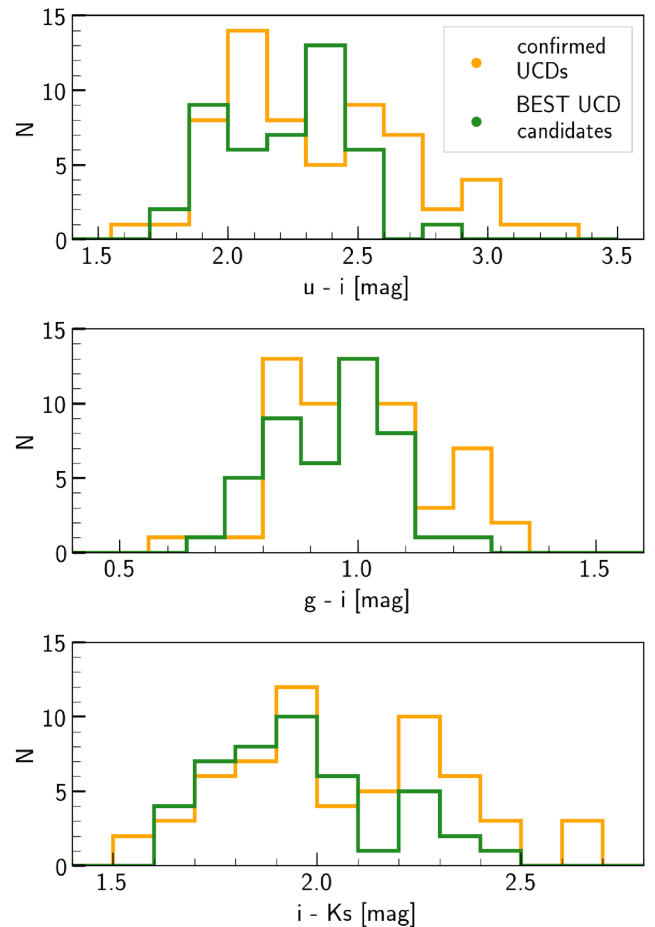


Figure 21. $u-i$, $g-i$, and $i-Ks$ colours of the BEST UCD candidates (in green) and the confirmed UCDs (in yellow).

the training-set. Next, we selected sources in ACSFCS with GC probability¹⁵ $p_{GC} > 0.75$ (GC probability from Jordán et al. 2015) and the projected distance to the host galaxy $R_{gal} > 30$ arcsec (3 kpc at the distance of the cluster). The limit on R_{gal} was made since the source extraction efficiency (for optical data, especially in r) drops significantly at small projected distances to galaxies. The largest compact object in ACSFCS has a half-light radius $r_h < 12$ pc. The GC selection criteria of the ACSFCS catalogue do not capture possible UCD/GCs larger than this size. As a consequence, some UCDs can be missed from this catalogue. One example is the recently discovered UCD around the low-mass elliptical galaxy FCC 47 (Fahrion et al. 2019). The outcome of the comparison between our candidates and ACSFCS is presented in the following:

(i) The sample of UCDs in ACSFCS contains 13 objects of which 7 objects are within our coverage (6 objects are around FCC21 in Fornax A, which is not covered in u band).

(ii) In our data, six of the seven objects are pre-selected (magnitude, size, and $u-i/i-Ks$ colours criteria). The missing object is in a low-signal-to-noise region in the Ks ; and its $i-Ks$ colour is not reliable.

(iii) Among the six pre-selected sources, all six are classified as UCD (four with UCD_SCORE = 100, one with UCD_SCORE = 98, and one with UCD_SCORE = 61). Although the matched sample is

¹⁵See Jordán et al. (2009) for more details.

small, hence might suffer for sampling effects, our source extraction and UCD selection when UCDs at $R_{\text{gal}} > 30$ arcsec are considered is 86 ± 15 per cent complete. Moreover, based on this sample, the recovery rate of the selection is 100 per cent, however, it is limited by the source extraction efficiency.

(iv) We also compared our classification outcome for objects with very low GC probability ($p_{\text{GC}} < 25$ per cent) which are very likely to be background galaxies. Our classifier did not pre-select these objects in the beginning and removed them from the list of possible UCDs.

Next, we extend our comparison with ACSFCS to GCs ($21 < m_g < 21.5$ mag) and investigate the accuracy of our classifier.

(i) The ACSFCS GC catalogue includes 32 GCs that are overlapping with our data set, within the magnitude range between $21 < m_g < 21.5$ mag and farther than $R_{\text{gal}} > 30$ arcsec from the host galaxy.

(ii) 28 out of 32 GCs are extracted from the data and 4 are missing. This means that we were able to extract 88 ± 6 per cent of the GCs in this magnitude range (source extraction efficiency). During the pre-selection step, 4 objects among 28 objects were excluded (not pre-selected).

(iii) At the end, among the 24 pre-selected sources, 23 are identified as GC with $\text{UCD_SCORE} > 95$ (20 with $\text{UCD_SCORE} = 100$) and 1 object is classified as a foreground star. This implies that the recovery rate of our selection is 72 ± 10 per cent. The misclassified object is located in the blue or red tail of the UCD sequence where miss-classification is most likely.

5.4 Optical versus optical/near-infrared UCD selection

Recently, Cantiello et al. (2020) published a catalogue of optically selected GCs in the Fornax cluster based on the FDS observations in u , g , r and i . Their catalogue includes 8971 UCD/GCs (1649 UCDs, based on the definition in that paper) and has a larger coverage than the covered area of our work since our data is limited to the data in J and K_s . Here, using the near-infrared photometry in J and K_s for the UCDs in Cantiello et al. (2020), we investigate the optical/near-infrared UCD/GC selection versus the optical selection (as in Cantiello et al. 2020).

(i) Among the 1649 GCs in Cantiello et al. (2020), 1168 objects have J and K_s photometry in our source catalogue, 35 and 1153 in the *KNOWN* and *UNKNOWN* catalogues.

(ii) Among these objects, after pre-selection (*step ii* in UCD/GC selection), 432 (around ~ 40 per cent of the initial number) are left. The $u-i-i-K_s$ colours of the other 736 objects (not-selected, around ~ 60 per cent of the initial number) show that, except for a few objects, the majority are foreground stars. This shows the well-known power of near-infrared photometry to exclude foreground stars. Among the 432 pre-selected sources, 35 are confirmed UCDs, 69 are foreground stars (spectroscopically confirmed), and 328 are unknown (no radial velocity available in our spectroscopic references).

(iii) Out of 328 *UNKNOWN* objects, our classifier has selected 7 as background galaxies, 195 as foreground star, and 126 as UCD (70 per cent with $\text{UCD_SCORE} > 80$).

(iv) In the end, 24 objects out of 126 ended up in our final BEST UCD catalogue. 26 of the BEST UCD candidates are not identified in Cantiello et al. (2020) of which 14 have $\text{FWHM}^* > 0.6$ and possibly are removed because of the adopted GC size criteria in Cantiello et al. (2020). Another 5 BEST candidates are objects brighter than $m_g = 19.5$ which is the upper magnitude limit of the UCD catalogue of Cantiello et al. (2020).

To summarize, out of 1168 sources in the UCD catalogue of Cantiello et al. (2020) with J and K_s data from our photometry, our classifier identified 126 as possible UCDs (not the BEST). Based on this result, while optical colours, including u band and size criteria are enough to identify background galaxies, to distinguish stars and UCDs efficiently, near-infrared colours are needed. Based on the $u-i-i-K_s$ colour-colour diagram, optically selected UCD/GC samples are at least 70 per cent contaminated by foreground stars (the contamination from background galaxies is small). Additionally, machine learning on the 5D colour-colour space rejects another 60 per cent on the remaining 30 per cent which leads to a total contamination rate of 90 per cent. Hence, as in Cantiello et al. (2020), a statistical decontamination from the background sources is required to characterize the properties of sources, if only optical data are used.

5.5 BEST UCD candidates in Gaia DR2 and EDR3

Once UCD candidates are identified, we checked if the sample is contaminated by *Gaia* stars. We selected *Gaia* sources with $\frac{\text{parallax}}{\text{parallax_error}} > 5.0$ as foreground (Milky Way) stars. Out of ~ 5000 *Gaia* stars with full photometric data in six filters, none were selected as UCD/GC and our classifier successfully excluded all these objects using their $u-i-i-K_s$ colours in the pre-selection step. However, this test is biased towards the population of the nearby stars which may not be a representative of the stars that can be confused with UCDs. Among the BEST UCD candidates, 41 (out of 44) have been identified in Gaia DR2 and they all have $\frac{\text{parallax}}{\text{parallax_error}} < 4.0$. However, in the most recent data release of *Gaia* (EDR3; Gaia Collaboration et al. 2021), the two brightest candidates, FDS-UCD-1 and FDS-UCD-2 have $\frac{\text{parallax}}{\text{parallax_error}} \sim 6.0$ (> 5.0) which shows that they are very likely foreground stars. While this ratio is a good indicator of foreground (Milky Way) stars, but it is not definitive. Therefore, we do not exclude these two candidates from our catalogue. The rest of the objects have $\frac{\text{parallax}}{\text{parallax_error}} < 5.0$ in Gaia EDR3.

Recently, Voggel et al. (2020) showed that UCD/GCs in the NGC 5128 group show larger BP/RP Flux excess values in Gaia DR2 and used this parameter to identify UCD/GCs around NGC 5128. As is described in the Gaia DR2 documentation,¹⁶ the derived magnitudes in BP and RP filters are measured within a window size of 3.5×2.1 arcsec and can suffer from nearby bright sources while G magnitudes are derived from accurate fitting of the objects within a small window and are less affected by the nearby sources. As a result, for blended objects or objects close to bright sources, the sum of the fluxes in BP and RP can be larger than the total flux in G . In the Gaia DR2, this ratio between two fluxes $(I_{BP} + I_{RP})/I_G$ is reported as BP/RP flux excess (Evans et al. 2018). The values of BP/RP flux excess closer to 1 indicate that the BP and RP magnitudes and colours of the corresponding object are reliable. For the unresolved foreground stars, the average BP/RP flux excess value is 1.20 mag. The larger BP/RP flux excess parameter can also happen when objects are resolved by *Gaia* (Voggel et al. 2020). Later, Liu et al. (2020) applied the same approach to investigate if their photometrically selected UCDs in the Virgo cluster show signs of being extended.

The Gaia DR2 data is very insensitive to the local background variations, however the background estimation has been improved in Gaia EDR3. The BP/RP flux excess is a complicated parameter and hard to interpret. Its behaviour depends on parameters such as colours of sources, magnitudes, variability, local background

¹⁶<https://gea.esac.esa.int/archive/documentation/GDR2/>

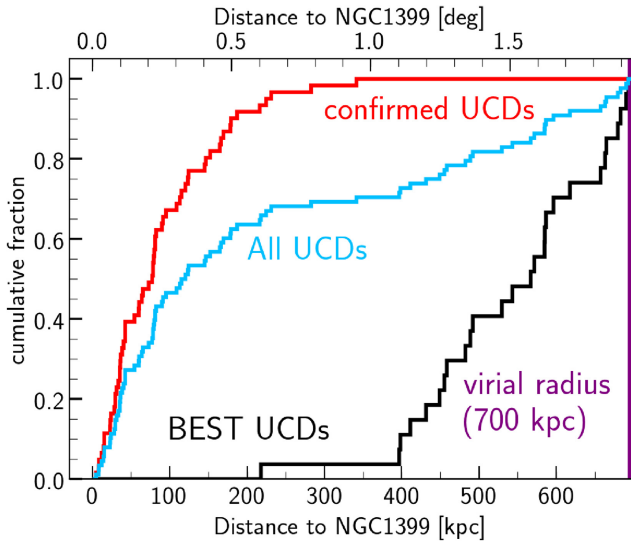


Figure 22. Cumulative radial distribution of the BEST UCD candidates (black), confirmed UCDs (red), and the combination of both (blue). Almost all of the BEST candidates are found outside the 360 kpc radius from NGC 1399. We assume that the spatial coverage of our data set up to the cluster’s virial radius is complete. Therefore, no completeness correction has been applied to the radial profile of the BEST UCD candidates.

variations, nearby bright objects, and processing problems. In case of the Fornax cluster, the confirmed UCDs are close to the bright galaxies and are located at the faint end of the *Gaia* detection limit. Therefore, while the approach discussed in Voggel et al. (2020) based on the *BP/RP* flux excess is useful for the nearby galaxy groups (such as NGC 5128 group as is studied in their work), for the Fornax cluster, it is too uncertain. Therefore, we did not apply this approach to our UCD selection.

6 DISCUSSION

6.1 The observed distribution of UCDs beyond the centre of the cluster

Earlier in this paper, we defined the terms ‘core’ and ‘outskirts’ of the Fornax cluster as the area within the virial radius of the cluster, inside and outside the central 1° (360 kpc, \sim half the virial radius) from NGC 1399. The majority of the identified BEST UCDs are located outside of the Fornax core, at a projected clustercentric distance larger than 360 kpc in a region that has not been explored before to this depth. The radial distribution of the confirmed UCDs, BEST UCDs, and the combined sample as representative of all the UCDs in the Fornax cluster is shown in Fig. 22. The UCD radial and projected distribution along with the distribution of Fornax cluster galaxies (Venhola et al. 2018) is investigated in Fig. 23. 70 per cent of the total UCDs within the virial radius are in the core of the cluster. UCDs in the outskirts are responsible for the remaining 30 per cent of the total UCD population. This fraction is in the same range as the fraction of massive galaxies in the outskirts (24 and 29 per cent for massive early-type and late-type galaxies, respectively) and is smaller than the fraction of dwarfs (53 and 81 per cent for early-type and late-type galaxies, respectively) and nucleated dwarfs (43 per cent). This implies that the UCDs are more likely to follow the massive galaxies and be connected to the high-density environments as it is expected. On the other hand, the projected distribution of the UCDs in Fig. 23 in the bottom panel

reveals overdensities of UCDs in the northern side and western side of the cluster, indicated by ‘N’ and ‘W’ in this figure. Both overdensities overlap with enhancements in population of dwarf galaxies in Fornax as shown in Venhola et al. (2019; fig. 3 in their paper) and Ordenes-Briceño et al. (2018a; fig. 3 in their paper). Thus, it suggests that a population of UCDs follow the dwarf galaxies and thus form in rather low-density, probably pre-processed group environments.

In the inner 360 kpc around NGC 1399, one object has been identified in a zone which is not covered in depth by previous surveys. This candidate (FDS-UCD-44) is not listed in any previous catalogue, it has a $UCD_SCORE = 92$ and is located inside the 60 per cent completeness contour in the $u-i/i-Ks$ diagram. It is unlikely that the spectroscopic surveys in the core of Fornax have missed many UCDs, since they are bright and point-like, and are easy to observe with 4- or 8-m class telescopes. However, not all UCD/GC candidates can be targeted in a multiobject approach with either slits or fibres and there will be always objects missed. If we assume that current spectroscopic surveys have covered almost all UCDs in the centre of the Fornax cluster, we find that among the BEST UCDs there is one false positive. This fact can be used to calculate an upper limit to the number of false positives in the Fornax cluster, and therefore a lower limit to the total number of UCDs in it. Given an average of one source within 360 kpc from NGC 1399, we estimate an average approximately six false-positives within the virial radius of the cluster (700 kpc), five outside the core (360 kpc). For this estimate, we took into account the projected number density of *UNKNOWN* sources in the core and in the outskirts. While the area in the outskirts is ~ 3 times larger than in the core, the number of *UNKNOWN* sources outside of the core is 10 times more than in the core. Considering the estimated background objects in the outskirts (10), the lower limit on the total number of the UCDs outside of the core is 33 ± 10 (within uncertainties consistent with the expected rate of true positives of the BEST UCD sample). Among them, 18 ± 7 are within the virial radius of the cluster. This lower limit implies that UCDs in the outskirts (within the virial radius of the cluster) are responsible for at least 23 ± 6 of the total UCDs within the virial radius.

While the majority of the BEST UCDs are outside the clusters’ core, a fraction of them is still found close to the massive galaxies in the cluster. Fig. 24 shows the projected distribution of the BEST UCDs and the brightest early-type/late-type Fornax cluster galaxies (Ferguson 1989; Venhola et al. 2018). The circles around each galaxy indicate a projected radius of 200 kpc from each galaxy. This value is comparable to the virial radius of the Milky Way ($R_{vir} \sim 200$ kpc; Dehnen, McLaughlin & Sachania 2006). 20 BEST UCDs are located outside any of these radii. Some of the BEST UCD candidates, if confirmed spectroscopically, can improve the record of the most isolated compact object from a major galaxy. The most isolated compact stellar system in the Local Group known up to now, MGC1 (Mackey et al. 2010) is located at de-projected (3D) distance of 200 kpc from M31. Outside of the Local Group, this record belongs to GC-2 in the M81 group. GC-2 has a de-projected distance of 400 kpc to M81 (Jang et al. 2012). All these objects are fainter than $M_g = -10.5$ (or $M_V = -11$ given the average colour $g-r \sim 0.7$ mag) and therefore fainter than the identified UCDs in the Fornax cluster.

If one assumes that the outskirts of galaxy clusters are formed from pre-processing in low-density environment such as groups, one can imagine that some of the nucleated group dEs/dIrrs (e.g. Georgiev et al. 2009) have been disrupted and left over the UCDs we see in the outskirts. In this case, a future follow-up spectroscopy and analysis of the radial velocities in a phase-space diagram can help to constrain the origin of these objects, whether they are in-falling systems or not.

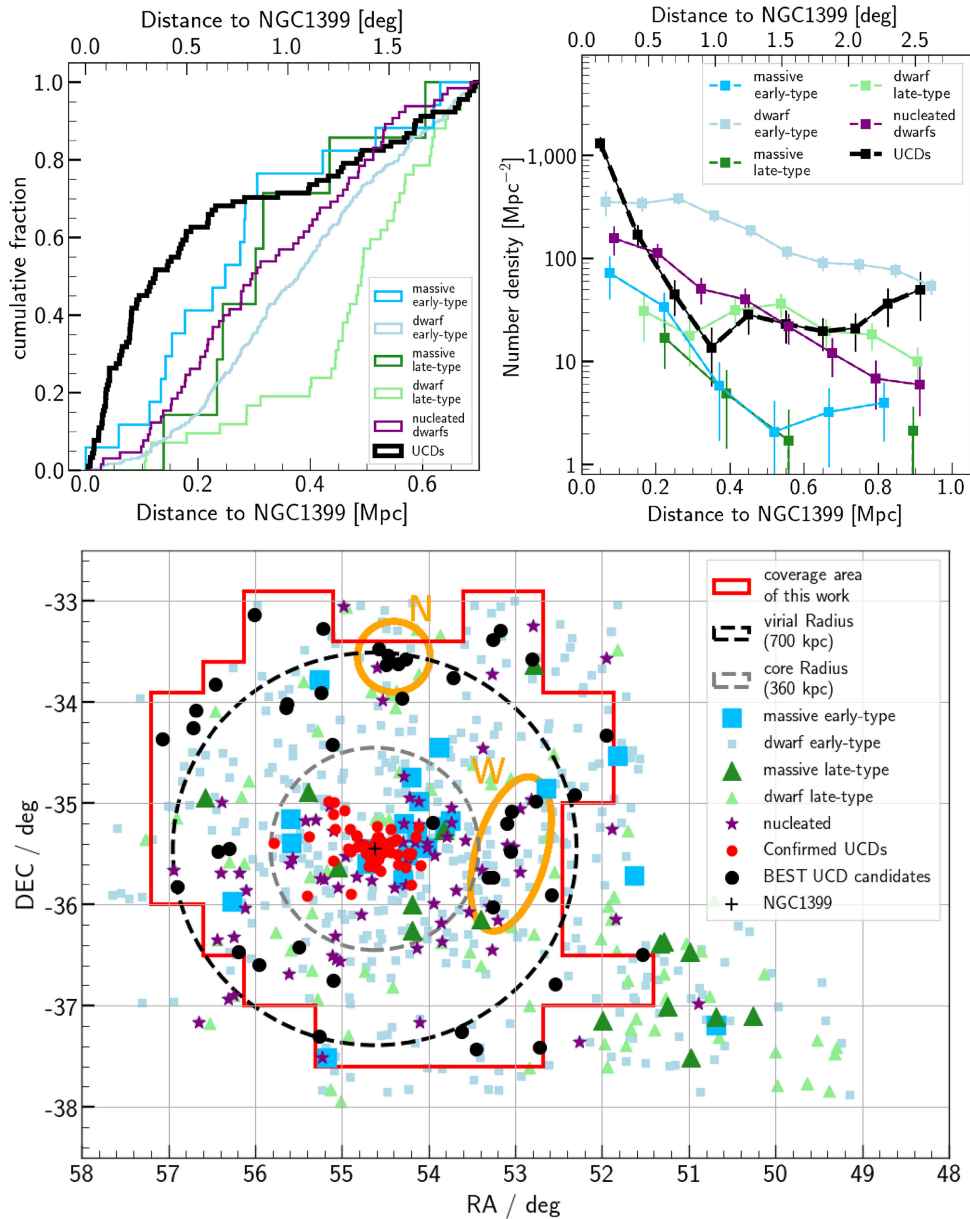


Figure 23. Top panels: Radial distribution of the UCDs in the Fornax cluster (black) and Fornax cluster galaxies of different types. The cumulative distributions (on the left-hand panel) are estimated up to the virial radius of the cluster (0.7 Mpc), where the data is spatially almost complete. The number density profiles are shown up to 1 Mpc, about 0.3 Mpc farther than the cluster’s virial radius, and are corrected for the spatial incompleteness.

Bottom panel: The projected distribution of the BEST UCDs and all the galaxies (massive, dwarf, nucleated) in the Fornax cluster. Massive/dwarf early-type and late-type galaxies are shown with larger/smaller blue squares and larger/smaller green triangles. Nucleated dwarf galaxies are indicated by purple stars. ‘N’ and ‘W’ indicate the two overdensities of UCDs in the northern side and western side of the cluster.

6.2 Rate of disruption of dwarf galaxies

It is expected that the disruption of dwarf galaxies in galaxy clusters impacts on the faint-end of the galaxy luminosity function. The observed number of UCDs in such environments can constrain the number of disrupted nucleated dwarf galaxies and the rate of dwarf galaxy disruption can be derived. Ferrarese et al. (2016) used the ratio between the numbers of UCDs (N_{UCD}) and NSCs (N_{NSC}) in the Virgo cluster to estimate the disruption rate of dwarf galaxies and correct the galaxy luminosity function based on this estimate. The authors derived $\frac{N_{\text{UCD}}}{N_{\text{NSC}}} = 2.8$ ($N_{\text{UCD}} = 92$ and $N_{\text{NSC}} = 33$) which implies that 70 per cent of all the infall haloes to the cluster did not

survive. In this calculation, two assumptions have been made: First, all the UCDs are stripped nuclei of dwarf galaxies, and secondly, the mass/luminosity of the UCD is the same as the mass/luminosity of the nucleus of the progenitor dwarf galaxy. The latter seems to be reasonable when the observed mass/luminosity of UCDs is compared to the predicated mass/luminosity of the nucleus derived from the observed mass of central SMBH (Graham 2020). However, the first assumption that all the observed UCDs are stripped nuclei is probably not correct. It is expected that the observed UCDs are a mixed bag of stripped nuclei and bright GCs, and that the ratio between stripped nuclei to GCs decreases with mass. Simulations of stripped nuclei

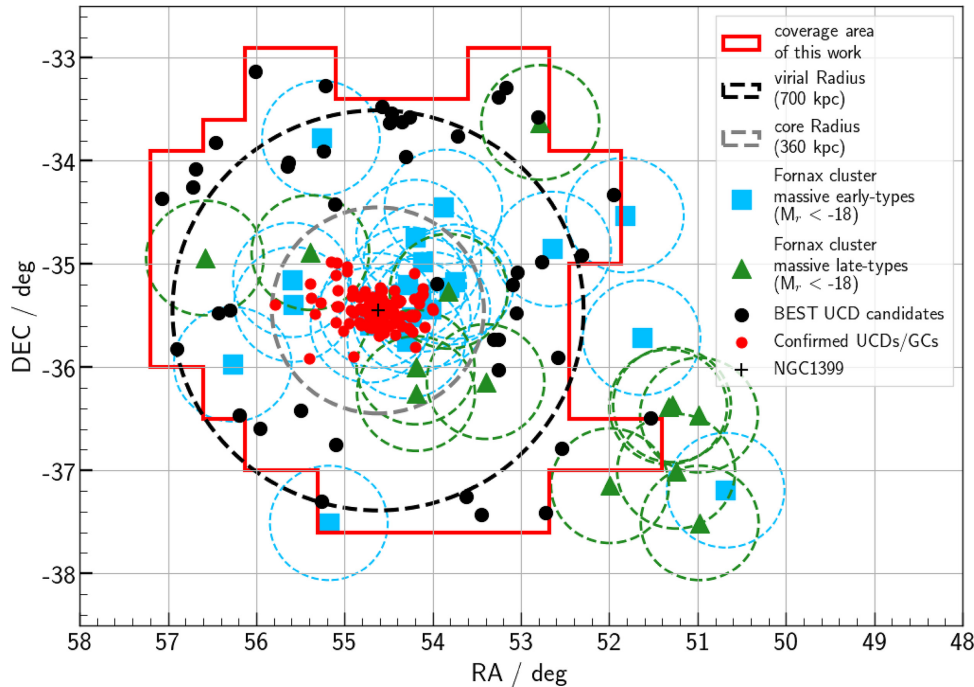


Figure 24. The projected distribution of all the BEST UCD candidates (black circles). All the confirmed UCD/GCs (shown in red) are located in the core of the cluster, within 360 kpc from NGC 1399 (grey dashed circle). Early-type and late-type Fornax cluster galaxies (Ferguson 1989; Venhola et al. 2018) brighter than $M_r = -18$ mag are shown with blue squares and green triangles. The circle around galaxies indicate a 200 kpc radius from them. Here, we use these circles to investigate the association of UCD candidates to the massive galaxies in the cluster. Almost half (46 per cent) of the BEST UCD candidates are further than 200 kpc from any massive galaxies.

(Pfeffer et al. 2014, 2016¹⁷) show that, while stripped nuclei are responsible for around half of the UCDs more massive than $10^7 M_\odot$, they are responsible for ~ 10 per cent of the UCDs more massive than $4 \times 10^6 M_\odot$ (same as our mass limit for UCDs). The ratio between N_{UCD} and N_{NSC} for the Fornax cluster galaxies is calculated in Voggel et al. (2019), in which only UCDs with observed $M_{\text{dyn}}/M_* > 1$ are considered (which are very likely to be a stripped nuclei). The authors found $\frac{N_{\text{UCD}}}{N_{\text{NSC}}} \sim 0.9$. This ratio implies that around half of the nucleated dwarfs are disrupted. Note that they only took into account the known UCDs in the central regions and used an older UCD/GC catalogue (also see section 9.2 of Neumayer, Seth & Böker (2020) for a discussion on UCD/NSC ratio).

The combination of the BEST UCD candidates (44 UCDs) in our work and the confirmed UCDs (61) provides a sample of 105 UCDs in the Fornax cluster of which 88 are within its virial radius. On the other hand, The FSDC (Venhola et al. 2018) contains 62 nucleated dwarf galaxies (fainter than $M_r = -18$ mag) within the virial radius of the cluster of which 13 have a nucleus brighter than $M_g = -10.5$ mag (same magnitude as the limiting magnitude of UCDs). For our calculations here, first we make the same assumptions as Ferrarese et al. (2016) and find $\frac{N_{\text{UCD}}}{N_{\text{NSC}}} = 6.8$ (lower limit 6.5 ± 0.6) which means that 87 per cent (lower limit 81 ± 8 per cent of the nucleated dwarfs (and dwarf galaxies in general) in the Fornax cluster have been disrupted to the present time. This rate is higher than the value estimated for the Virgo cluster (70 per cent; Ferrarese et al. 2016).

As was discussed above, not all the observed UCDs are stripped nuclei. Therefore, next, we assume that only 10 per cent of the UCDs are a stripped nucleus and the rest are GCs (Pfeffer et al. 2014, 2016). This assumption is not consistent with how Ferrarese et al. (2016) did this. We find a UCD to NSC ratio of $\frac{N_{\text{UCD}}}{N_{\text{NSC}}} = 0.68$ for the Fornax cluster. This implies that about 40 per cent of the nucleated dwarfs have been disrupted. $\frac{N_{\text{UCD}}}{N_{\text{NSC}}}$ for the core and outskirts of the cluster is 0.88 and 0.43 which implies a rate of 47 and 30 per cent of the nucleated dwarfs in the core and in the outskirts are stripped. Note that these values depend very strongly on our assumption for the ratio between the number of stripped nuclei to the total number of observed UCDs.

6.3 Fornax versus Virgo

Recently Liu et al. (2020) published a catalogue of UCD candidates in the Virgo cluster. The majority of these objects are selected using optical/near-infrared imaging data of the Next-Generation Virgo Cluster Survey (NGVS) in $u, g, r, i,$ and K_s from the UKIRT Infrared Deep Sky Survey (UKIDSS). In their selection, they defined UCDs as objects brighter than $m_g = 21.5$ ($M_g = -9.59$ at the distance of the Virgo cluster). Therefore, their magnitude limit for defining UCDs is ~ 0.9 fainter than our adopted magnitude limit of UCDs. Additionally, the authors define a lower limit on sizes of objects and select UCDs as objects with a half-light radius (effective radius) larger than $r_h = 10$ pc. In Liu et al. (2020), the authors published a catalogue of UCD candidates in the Virgo cluster. This catalogue provides an opportunity to compare the results of the Virgo cluster with our results of the Fornax cluster. Given the rich populations of massive and bright galaxies in the Virgo cluster, it is not possible to identify the possible free-floating/remote UCDs, far from the massive galaxies.

¹⁷Note that simulation of stripped nuclei in Pfeffer et al. (2014, 2016) has limitations in the known nucleation fraction and missed the low-surface brightness (LSB) population.

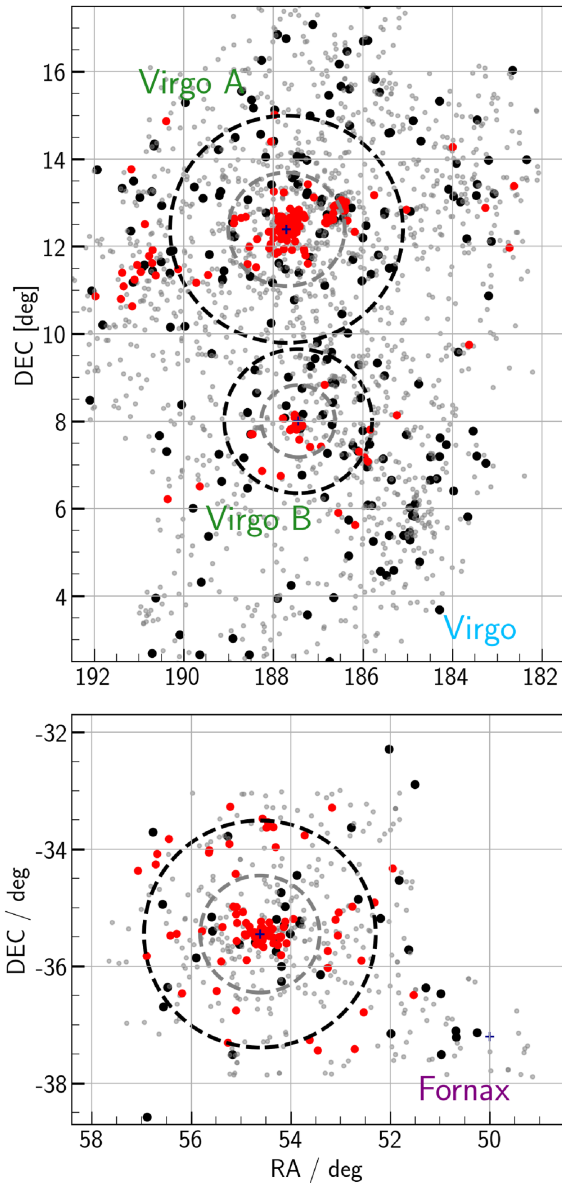


Figure 25. UCDs in Virgo (top panel) versus Fornax (bottom panel). UCDs are shown in red. Galaxies brighter than $M_B = -17$ are shown in black and are from VCC (Binggeli, Sandage & Tammann 1985) and FCC (Ferguson 1989) catalogues. Grey points indicate galaxies, fainter than $M_B = -17$ from VCC (Binggeli et al. 1985) and FDS (Venhola et al. 2018) catalogues. The virial radius (black dashed circles) and inner region (half the virial radius, grey dashed circles) for Virgo A cluster (around M87) and Virgo B cluster (around M49) and Fornax cluster (around NGC 1399) are indicated by black and grey circles.

The distribution of the UCDs and the brightest galaxies ($M_B > -17$) in Virgo and Fornax are shown in Fig. 25. For the UCDs in the Virgo cluster (top panel), we used the published catalogue of Liu et al. (2020) and selected spectroscopically (confirmed) and photometrically (optical/near-infrared) identified UCDs brighter than $m_g = 20.6$ mag ($M_g = -10.5$ at the distance of the Virgo cluster). Of 612 UCDs in their main UCD sample, our selection criterion gives 109 and 24 UCDs within the virial radius of Virgo A and B clusters. For the UCDs in the Fornax cluster (bottom panel), we used the spectroscopically (confirmed) UCDs and the BEST UCD candidates. Additionally, for the BEST UCDs, we apply the size

limit $r_h > 10$ pc to be consistent with the UCD sample of the Virgo cluster. This selection led to 87 UCDs within its virial radius. For Virgo A, B, and Fornax, we counted 13, 7, and 26 UCDs outside of the cluster's core (grey-dashed circle) corresponding to 12 ± 4 , 29 ± 11 , and 30 ± 6 per cent of the total UCDs within the virial radius. This implies that Fornax and Virgo B have more UCDs in their outskirts, almost twice more than Virgo A. Spectroscopic observations are needed to investigate the significance and the origin of this finding.

7 CONCLUSION AND SUMMARY

Our current knowledge of UCDs comes from limited and biased samples of confirmed UCDs. The known UCDs were found through spectroscopic surveys around massive galaxies or in the cores of galaxy clusters/groups. Therefore, there is not much known about UCDs (or similar objects) which might be located in the outskirts of galaxy clusters. In the Fornax cluster, all of the confirmed UCDs are located within half the virial radius of the cluster.

Using data from the FDS and the near-infrared observations with VISTA/VIRCAM (six filters in total: u , g , r , i , J , and K_s), we identified, for the first time, UCD candidates in the outskirts of the Fornax cluster. However, follow-up spectroscopy is necessary to confirm the membership of the candidates.

We prepared two catalogues containing the optical and near-infrared magnitudes, colours, and sizes of the sources with and without spectroscopic data (radial velocities) based on our spectroscopic references (Wittmann et al. 2016; Pota et al. 2018; Maddox et al. 2019). Then, we pre-selected UCD/GCs based on their magnitude, size, and $u-i-i-K_s$ colours and used the objects with available spectroscopic data (radial velocities) in the direction of the Fornax cluster as a reference-set (training-set) for a supervised machine learning method, evaluating a 5D colour space. As the result, all the pre-selected sources were classified into three classes: (i) foreground stars, (ii) UCD/GCs, and (iii) background galaxies. Furthermore, we examined our classification with the *HST*/ACS observations of the Fornax cluster and Gaia DR2.

Among the ~ 220 initially selected UCD candidates, we identified 44 BEST UCD candidates, which have a higher probability of being real UCDs. We also estimated that combining the near-infrared observations with the optical helps to remove 60–90 per cent of the contamination from foreground stars for the optically selected UCD. Based on our analysis on UCDs in the Fornax cluster in this paper, we conclude:

(i) UCDs outside of the core and within the virial radius of the cluster are responsible for 30 per cent (lower limit 23 ± 6 per cent) of the UCDs within the virial radius, more or less the same as the most massive galaxies in the cluster (early-type or late-type).

(ii) We identified two overdensities outside the core of the cluster in the northern and western sides, overlapping with the enhancements in the densities of the Fornax cluster dwarf galaxies (as is shown in the literature). This implies that a population of UCDs follows the dwarf galaxies in the cluster and may form in low-density, pre-processed group environments.

(iii) Almost half (46 per cent) of the identified UCDs in the outskirts of the Fornax cluster are further away than 200 kpc from any Fornax galaxy brighter than $M_r = -18$ mag. The exact origin of this population is not clear since they can be UCDs associated to massive galaxies with an highly elongated orbit or UCDs formed in low-density environments (such as galaxy groups) and represent in-falling UCD populations. Future spectroscopic follow-ups to measure their radial velocities can constrain their origin.

(iv) We estimated that 87 per cent of the nucleated dwarfs (and dwarfs in general) in the Fornax cluster are stripped. This rate, compared to the estimated rate in the Virgo cluster (70 per cent; Ferrarese et al. 2016) shows a higher disruption rate in Fornax versus Virgo. While these rates are strongly depend on the assumption on the fraction of stripped nuclei among the present day UCDs (100 per cent for the expressed values here), the relatively higher disruption rate in the Fornax cluster is independent of this assumption.

(v) A comparison between the UCD populations in Virgo (Liu et al. 2020) and Fornax (this work) shows that the Fornax cluster and Virgo B cluster (M49 sub-cluster) have relatively more UCDs in their outskirts (within their virial radius) than Virgo A (M87 sub-cluster).

Until now, most studies focused on compact sources around galaxies in the cores of galaxy clusters. This is mainly because of the lack of proper observations in the clusters' outskirts. Recently, Liu et al. (2020) published a catalogue of UCD candidates in the Virgo cluster, including intra-cluster UCD candidates. In the case of Fornax cluster, in this work, we publish a catalogue of UCD candidates up to the virial radius of the cluster and a bit further. We investigated the properties of these candidates and their distribution which hints at a population of UCDs in low-density environments. However, the lack of realistic simulations in the outskirts of the clusters makes it very tricky to interpret our results. Upcoming cosmological simulation of GCs such as E-MOSAICS (Pfeffer et al. 2018; Kruijssen et al. 2019) can provide a framework to compare our results with the model predictions. Follow-up spectroscopy and radial velocity measurements in the future can confirm/reject the membership of the candidates and investigate the origin of the UCDs beyond the centre of the Fornax cluster.

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(van der Walt, Colbert & Varoquaux 2011), SCIPY (Virtanen et al. 2020), ASTROPY (Astropy Collaboration et al. 2018), and SCIKIT-LEARN (Pedregosa et al. 2011).

DATA AVAILABILITY

The final mosaics of the optical data of FDS, the final mosaics of the near-infrared data of VHS, and the raw frames of the VISTA/VIRCAM observations are available through ESO Science Archive Facility. The final mosaics of the *Ks* data are not available yet, however, they will be available soon. The data bases, software, packages, and libraries that are used throughout this paper, SIMBAD, VIZIER, DS9, ALADIN, SEXTRACTOR, SWARP, ASTROPY, SCIKIT-LEARN are publicly available. The catalogue of the UCD and GC candidates in the Fornax cluster generated in this research is available in the article and in its online supplementary material.

REFERENCES

- Afanasiev A. V. et al., 2018, *MNRAS*, 477, 4856
 Ahn C. P. et al., 2017, *ApJ*, 839, 72
 Ahn C. P. et al., 2018, *ApJ*, 858, 102
 Anders P., Gieles M., de Grijs R., 2006, *A&A*, 451, 375
 Angora G. et al., 2019, *MNRAS*, 490, 4080
 Astropy Collaboration et al., 2018, *AJ*, 156, 123
 Baumgardt H., Hilker M., 2018, *MNRAS*, 478, 1520
 Baumgardt H., Makino J., Hut P., McMillan S., Portegies Zwart S., 2003, *ApJ*, 589, L25
 Begeman K., Belikov A. N., Boxhoorn D. R., Valentijn E. A., 2013, *Exp. Astron.*, 35, 1
 Bekki K., Couch W. J., Drinkwater M. J., Shioya Y., 2003, *MNRAS*, 344, 399
 Bergond G. et al., 2007, *A&A*, 464, L21
 Bertin E., Arnouts S., 1996, *A&AS*, 117, 393
 Bertin E., Mellier Y., Radovich M., Missonnier G., Didelon P., Morin B., 2002, in Bohlender D. A., Durand D., Handley T. H., eds, *Astronomical Data Analysis Software and Systems XI*, Vol. 281, The TERAPIX Pipeline. Astron. Soc. Pac., San Francisco, p. 228
 Binggeli B., Sandage A., Tammann G. A., 1985, *AJ*, 90, 1681
 Blakeslee J. P. et al., 2009, *ApJ*, 694, 556
 Boch T., Fernique P., 2014, in Manset N., Forshay P., eds, *ASP Conf. Ser. Vol. 485*, *Astronomical Data Analysis Software and Systems XXIII*. Astron. Soc. Pac., San Francisco, p. 277
 Bonnarel F. et al., 2000, *A&AS*, 143, 33
 Brodie J. P., Romanowsky A. J., Strader J., Forbes D. A., 2011, *AJ*, 142, 199
 Cantiello M., Grado A., Rejkuba M., Arnaboldi M., Capaccioli M., Greggio L., Iodice E., Limatola L., 2018, *A&A*, 611, A21
 Cantiello M. et al., 2020, *A&A*, 639, 23
 Conroy C., Loeb A., Spergel D. N., 2011, *ApJ*, 741, 72
 Cover T., Hart P., 1967, *IEEE Trans. Inf. Theory*, 13, 21
 Covey K. R. et al., 2007, *AJ*, 134, 2398
 D'Souza R., Rix H.-W., 2013, *MNRAS*, 429, 1887
 Da Rocha C., Mieske S., Georgiev I. Y., Hilker M., Ziegler B. L., Mendes de Oliveira C., 2011, *A&A*, 525, A86
 Dehnen W., McLaughlin D. E., Sachania J., 2006, *MNRAS*, 369, 1688
 den Brok M. et al., 2014, *MNRAS*, 445, 2385
 Drinkwater M. J., Jones J. B., Gregg M. D., Phillips S., 2000a, *PASA*, 17, 227
 Drinkwater M. J. et al., 2000b, *A&A*, 355, 900
 Drinkwater M. J., Gregg M. D., Colless M., 2001, *ApJ*, 548, L139
 Ochsenbein F., Bauer P., Marcout J., 2000, *A&AS*, 143, 23
 Evans D. W. et al., 2018, *A&A*, 616, A4
 Evstigneeva E. A. et al., 2008, *AJ*, 136, 461
 Fahrion K. et al., 2019, *A&A*, 625, A50
 Fahrion K. et al., 2020a, *A&A*, 634, A53
 Fahrion K. et al., 2020b, *A&A*, 637, A26
 Fellhauer M., Kroupa P., 2002, *MNRAS*, 330, 642

- Ferguson H. C., 1989, *AJ*, 98, 367
- Ferrarese L. et al., 2016, *ApJ*, 824, 10
- Firth P., Drinkwater M. J., Evstigneeva E. A., Gregg M. D., Karick A. M., Jones J. B., Phillipps S., 2007, *MNRAS*, 382, 1342
- Firth P., Drinkwater M. J., Karick A. M., 2008, *MNRAS*, 389, 1539
- Firth P., Evstigneeva E. A., Drinkwater M. J., 2009, *MNRAS*, 394, 1801
- Forbes D. A., Norris M. A., Strader J., Romanowsky A. J., Pota V., Kannappan S. J., Brodie J. P., Huxor A., 2014, *MNRAS*, 444, 2993
- Forbes D. A., Ferré-Mateu A., Durré M., Brodie J. P., Romanowsky A. J., 2020, *MNRAS*, 497, 765
- Francis K. J., Drinkwater M. J., Chilingarian I. V., Bolt A. M., Firth P., 2012, *MNRAS*, 425, 325
- Gaia Collaboration et al., 2018, *A&A*, 616, A1
- Gaia Collaboration et al., 2021, *A&A*, 649, A1
- Georgiev I. Y., Hilker M., Puzia T. H., Goudfrooij P., Baumgardt H., 2009, *MNRAS*, 396, 1075
- González-Lópezlira R. A. et al., 2019, *ApJ*, 876, 39
- Graham A. W., 2020, *MNRAS*, 492, 3263
- Gregg M. D. et al., 2009, *AJ*, 137, 498
- Haghi H., Safaei G., Hasani Zonoozi A., Kroupa P., 2020, *ApJ*, 904, 17
- Hilker M., 2006, preprint ([arXiv:astro-ph/0605447](https://arxiv.org/abs/astro-ph/0605447))
- Hilker M., Infante L., Vieira G., Kissler-Patig M., Richtler T., 1999, *A&AS*, 134, 75
- Ibata R., Nipoti C., Sollima A., Bellazzini M., Chapman S. C., Dalessandro E., 2013, *MNRAS*, 428, 3648
- Iodice E. et al., 2016, *ApJ*, 820, 42
- Iodice E. et al., 2017, *ApJ*, 839, 21
- Iodice E. et al., 2019, *A&A*, 627, A136
- Ivezić Ž. et al., 2004, *Astron. Nachr.*, 325, 583
- Jang I. S., Lim S., Park H. S., Lee M. G., 2012, *ApJ*, 751, L19
- Janz J. et al., 2016, *MNRAS*, 456, 617
- Janz J., Forbes D. A., Norris M. A., Strader J., Penny S. J., Fagioli M., Romanowsky A. J., 2015, *MNRAS*, 449, 1716
- Jerjen H., 2003, *A&A*, 398, 63
- Jester S. et al., 2005, *AJ*, 130, 873
- Johnston E. J. et al., 2020, *MNRAS*, 495, 2247
- Jordán A. et al., 2004, *ApJS*, 154, 509
- Jordán A. et al., 2007, *ApJS*, 169, 213
- Jordán A. et al., 2009, *ApJS*, 180, 54
- Jordán A., Peng E. W., Blakeslee J. P., Côté P., Eyheramendy S., Ferrarese L., 2015, *ApJS*, 221, 13
- Joye W. A., Mandel E., 2003, in Payne H. E., Jedrzejewski R. I., Hook R. N., eds, ASP Conf. Ser. Vol. 295, Astronomical Data Analysis Software and Systems XII. Astron. Soc. Pac., San Francisco, p. 489
- King I. R., 1966, *AJ*, 71, 276
- Kron R. G., 1980, *ApJS*, 43, 305
- Kroupa P., 2020, in Bragaglia A., Davies M., Sills A., Vesperini E., eds, Proc. IAU Symp. 351, Star Clusters: From the Milky Way to the Early Universe. Kluwer, Dordrecht, p. 117
- Kruijssen J. M. D., 2008, *A&A*, 486, L21
- Kruijssen J. M. D., Pfeffer J. L., Crain R. A., Bastian N., 2019, *MNRAS*, 486, 3134
- Kuijken K. et al., 2019, *A&A*, 625, A2
- Liu C. et al., 2015, *ApJ*, 812, 34
- Liu C. et al., 2020, *ApJS*, 250, 28
- Mackey A. D. et al., 2010, *MNRAS*, 401, 533
- Maddox N., Serra P., Venhola A., Peletier R., Loubser I., Iodice E., 2019, *MNRAS*, 490, 1666
- Mahani H., Zonoozi A. H., Haghi H., Jeřábková T., Kroupa P., Mieske S., 2021, *MNRAS*, 502, 5185
- Mashchenko S., Sills A., 2005, *ApJ*, 619, 243
- Mayes R. J., Drinkwater M. J., Pfeffer J., Baumgardt H., Liu C., Ferrarese L., Côté P., Peng E. W., 2021, *MNRAS*, 501, 1852
- McFarland J. P., Verdoes-Kleijn G., Sikkema G., Helmich E. M., Boxhoorn D. R., Valentijn E. A., 2013, *Exp. Astron.*, 35, 45
- McMahon R. G., Banerji M., Gonzalez E., Kuposov S. E., Bejar V. J., Lodieu N., Rebolo R., VHS Collaboration, 2013, *The Messenger*, 154, 35
- Mieske S. et al., 2008, *A&A*, 487, 921
- Mieske S., Hilker M., Infante L., 2002, *A&A*, 383, 823
- Mieske S., Hilker M., Infante L., 2004, *A&A*, 418, 445
- Mieske S., Frank M. J., Baumgardt H., Lützgendorf N., Neumayer N., Hilker M., 2013, *A&A*, 558, A14
- Misgeld I., Hilker M., 2011, *MNRAS*, 414, 3699
- Muñoz R. P. et al., 2014, *ApJS*, 210, 4
- Neumayer N., Seth A., Böker T., 2020, *A&A Rev.*, 28, 4
- Norris M. A. et al., 2014, *MNRAS*, 443, 1151
- Norris M. A., Escudero C. G., Faifer F. R., Kannappan S. J., Forte J. C., van den Bosch R. C. E., 2015, *MNRAS*, 451, 3615
- Ordenes-Briceño Y. et al., 2018a, *ApJ*, 859, 52
- Ordenes-Briceño Y. et al., 2018b, *ApJ*, 860, 4
- Pedregosa F. et al., 2011, *J. Mach. Learn. Res.*, 12, 2825
- Peletier R. et al., 2020, preprint ([arXiv:2008.12633](https://arxiv.org/abs/2008.12633))
- Pfeffer J., Griffen B. F., Baumgardt H., Hilker M., 2014, *MNRAS*, 444, 3670
- Pfeffer J., Hilker M., Baumgardt H., Griffen B. F., 2016, *MNRAS*, 458, 2492
- Pfeffer J., Kruijssen J. M. D., Crain R. A., Bastian N., 2018, *MNRAS*, 475, 4309
- Phillipps S., Drinkwater M. J., Gregg M. D., Jones J. B., 2001, *ApJ*, 560, 201
- Pota V. et al., 2018, *MNRAS*, 481, 1744
- Powalka M. et al., 2017, *ApJ*, 844, 104
- Prole D. J. et al., 2019, *MNRAS*, 484, 4865
- Raj M. A. et al., 2020, *A&A*, 640, A137
- Richtler T., Dirsch B., Larsen S., Hilker M., Infante L., 2005, *A&A*, 439, 533
- Rijsbergen C. J. V., 1979, *Information Retrieval*, 2nd edn. Butterworth-Heinemann, Oxford
- Schuberth Y., Richtler T., Hilker M., Dirsch B., Bassino L. P., Romanowsky A. J., Infante L., 2010, *A&A*, 513, A52
- Schweizer F., Seitzer P., Whitmore B. C., Kelson D. D., Villanueva E. V., 2018, *ApJ*, 853, 54
- Seth A. C. et al., 2014, *Nature*, 513, 398
- Sharina M. E., Puzia T. H., Makarov D. I., 2005, *A&A*, 442, 85
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Spavone M. et al., 2020, *A&A*, 639, A14
- Taylor M. A., Muñoz R. P., Puzia T. H., Mieske S., Eigenthaler P., Bovill M. S., 2016, preprint ([arXiv:1608.07285](https://arxiv.org/abs/1608.07285))
- van der Walt S., Colbert S. C., Varoquaux G., 2011, *Comput. Sci. Eng.*, 13, 22
- Vazdekis A., Sánchez-Blázquez P., Falcón-Barroso J., Cenarro A. J., Beasley M. A., Cardiel N., Gorgas J., Peletier R. F., 2010, *MNRAS*, 404, 1639
- Vazdekis A., Koleva M., Ricciardelli E., Röck B., Falcón-Barroso J., 2016, *MNRAS*, 463, 3409
- Venhola A. et al., 2017, *A&A*, 608, A142
- Venhola A. et al., 2018, *A&A*, 620, A165
- Venhola A. et al., 2019, *A&A*, 625, A143
- Villaume A., Brodie J., Conroy C., Romanowsky A. J., van Dokkum P., 2017, *ApJ*, 850, L14
- Virtanen P. et al., 2020, *Nature Methods*, 17, 261
- Voggel K., Hilker M., Baumgardt H., Collins M. L. M., Grebel E. K., Husemann B., Richtler T., Frank M. J., 2016a, *MNRAS*, 460, 3384
- Voggel K., Hilker M., Richtler T., 2016b, *A&A*, 586, A102
- Voggel K. T., Seth A. C., Baumgardt H., Mieske S., Pfeffer J., Rasskazov A., 2019, *ApJ*, 871, 159
- Voggel K. T., Seth A. C., Sand D. J., Hughes A., Strader J., Crnojevic D., Caldwell N., 2020, *ApJ*, 899, 140
- Wenger M. et al., 2000, *A&AS*, 143, 9
- Wittmann C., Lisker T., Pasquali A., Hilker M., Grebel E. K., 2016, *MNRAS*, 459, 4450
- Zhang H.-X. et al., 2018, *ApJ*, 858, 37

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APPENDIX: CATALOGUES

Table A1. Catalogue of the spectroscopically confirmed UCD/GC (sample table, the full table is available online). Columns from left- to right-hand represent RA (a), Declination (b), magnitude in u , g , r , i , J , and Ks (c–h), measured half-light (effective) radius (i). Quoted errors are statistical ones, systematic errors are not included.

RA (hms) (a)	Dec. (dms) (b)	u (AB mag) (c)	g (AB mag) (d)	r (AB mag) (e)	i (AB mag) (f)	J (Vega mag) (g)	Ks (Vega mag) (h)	r_h (pc) (i)
03 ^h 38 ^m 54 ^s 0	−35° 33′ 33″.01	20.78 ± 0.014	19.28 ± 0.010	18.28 ± 0.016	17.84 ± 0.029	16.19 ± 0.010	15.27 ± 0.006	139.0
03 ^h 38 ^m 5 ^s 04	−35° 24′ 09″.48	20.87 ± 0.013	19.37 ± 0.015	18.59 ± 0.006	18.28 ± 0.005	16.74 ± 0.014	16.03 ± 0.012	43.6
03 ^h 39 ^m 35 ^s 9	−35° 28′ 24″.57	21.14 ± 0.017	19.74 ± 0.010	19.16 ± 0.016	18.71 ± 0.029	17.30 ± 0.018	16.55 ± 0.019	43.9
03 ^h 38 ^m 6 ^s 29	−35° 28′ 58″.65	21.29 ± 0.019	19.86 ± 0.016	19.14 ± 0.006	18.74 ± 0.006	17.21 ± 0.018	16.53 ± 0.019	37.5
03 ^h 38 ^m 6 ^s 29	−35° 28′ 58″.65	21.29 ± 0.019	19.86 ± 0.016	19.14 ± 0.006	18.74 ± 0.006	17.21 ± 0.018	16.53 ± 0.019	37.5
03 ^h 37 ^m 3 ^s 22	−35° 38′ 04″.54	21.46 ± 0.024	19.96 ± 0.039	19.19 ± 0.035	18.80 ± 0.016	17.22 ± 0.018	16.47 ± 0.017	35.0
03 ^h 38 ^m 10 ^s 3	−35° 24′ 05″.79	21.45 ± 0.022	20.01 ± 0.016	19.26 ± 0.006	18.97 ± 0.006	17.63 ± 0.022	16.84 ± 0.025	15.3
03 ^h 39 ^m 52 ^s 5	−35° 04′ 24″.00	21.23 ± 0.022	20.02 ± 0.030	19.45 ± 0.088	19.17 ± 0.040	17.71 ± 0.110	17.13 ± 0.032	38.0
03 ^h 38 ^m 23 ^s 7	−35° 13′ 49″.48	21.31 ± 0.020	20.14 ± 0.002	19.56 ± 0.015	19.33 ± 0.004	18.08 ± 0.028	17.42 ± 0.042	18.3
03 ^h 37 ^m 43 ^s 5	−35° 22′ 51″.38	21.40 ± 0.022	20.20 ± 0.006	19.60 ± 0.023	19.31 ± 0.006	18.10 ± 0.030	17.32 ± 0.039	25.5
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Table A2. Catalogue of the spectroscopic foreground stars (sample table, the full table is available online). Columns from left- to right-hand represent RA (a), Declination (b), magnitude in u , g , r , i , J , and Ks (c–h). Quoted errors are statistical ones, systematic errors are not included.

RA (hms) (a)	Dec. (dms) (b)	u (AB mag) (c)	g (AB mag) (d)	r (AB mag) (e)	i (AB mag) (f)	J (Vega mag) (g)	Ks (Vega mag) (h)
03 ^h 31 ^m 38 ^s 5	−35° 53′ 13″.4	16.95 ± 0.004	15.89 ± 0.027	15.90 ± 0.027	16.10 ± 0.024	15.31 ± 0.007	15.54 ± 0.017
03 ^h 32 ^m 43 ^s 2	−35° 48′ 24″.1	17.29 ± 0.011	16.30 ± 0.032	16.25 ± 0.028	16.22 ± 0.027	15.54 ± 0.008	15.45 ± 0.012
03 ^h 36 ^m 18 ^s 8	−35° 14′ 58″.6	16.92 ± 0.010	16.34 ± 0.020	16.15 ± 0.033	16.15 ± 0.023	15.33 ± 0.008	15.17 ± 0.006
03 ^h 38 ^m 31 ^s 0	−35° 16′ 38″.9	17.31 ± 0.003	16.37 ± 0.001	16.45 ± 0.015	16.59 ± 0.002	16.00 ± 0.010	16.15 ± 0.013
03 ^h 37 ^m 14 ^s 3	−35° 17′ 41″.0	17.23 ± 0.007	16.38 ± 0.005	15.96 ± 0.023	15.94 ± 0.004	14.78 ± 0.010	14.51 ± 0.003
03 ^h 38 ^m 18 ^s 0	−34° 44′ 38″.8	17.48 ± 0.022	16.42 ± 0.008	16.06 ± 0.007	15.96 ± 0.007	15.11 ± 0.022	14.78 ± 0.009
03 ^h 40 ^m 21 ^s 8	−36° 10′ 40″.6	17.72 ± 0.030	16.42 ± 0.031	16.05 ± 0.025	15.96 ± 0.033	15.01 ± 0.012	14.72 ± 0.006
03 ^h 28 ^m 54 ^s 3	−34° 55′ 21″.1	17.14 ± 0.026	16.45 ± 0.048	16.04 ± 0.032	15.96 ± 0.042	15.00 ± 0.015	14.65 ± 0.008
03 ^h 37 ^m 42 ^s 4	−35° 29′ 54″.7	17.33 ± 0.006	16.46 ± 0.020	16.07 ± 0.017	15.99 ± 0.004	15.19 ± 0.011	14.91 ± 0.004
03 ^h 30 ^m 00 ^s 3	−35° 29′ 56″.8	17.22 ± 0.005	16.47 ± 0.004	16.18 ± 0.027	16.00 ± 0.004	15.13 ± 0.011	14.79 ± 0.008
–	–	–	–	–	–	–	–

Table A3. Catalogue of the spectroscopic background galaxies (sample table, the full table is available online). Columns from left- to right-hand represent RA (a), Declination (b), magnitude in u , g , r , i , J , and Ks (c–h). Quoted errors are statistical ones, systematic errors are not included.

RA (hms) (a)	Dec. (dms) (b)	u (AB mag) (c)	g (AB mag) (d)	r (AB mag) (e)	i (AB mag) (f)	J (Vega mag) (g)	Ks (Vega mag) (h)
03 ^h 38 ^m 46 ^s 8	−35° 47′ 17″.4	17.47 ± 0.002	17.53 ± 0.030	17.37 ± 0.013	17.13 ± 0.014	16.06 ± 0.012	15.11 ± 0.005
03 ^h 35 ^m 40 ^s 1	−34° 07′ 49″.3	17.77 ± 0.021	17.56 ± 0.009	17.78 ± 0.031	17.59 ± 0.013	16.76 ± 0.023	15.39 ± 0.018
03 ^h 35 ^m 18 ^s 5	−34° 44′ 03″.5	19.94 ± 0.013	18.16 ± 0.008	17.41 ± 0.022	16.85 ± 0.014	15.12 ± 0.008	14.19 ± 0.009
03 ^h 30 ^m 59 ^s 3	−34° 14′ 57″.0	18.93 ± 0.009	18.17 ± 0.007	17.56 ± 0.003	17.23 ± 0.019	15.50 ± 0.010	14.64 ± 0.007
03 ^h 34 ^m 13 ^s 4	−33° 36′ 26″.2	18.77 ± 0.006	18.25 ± 0.002	17.72 ± 0.004	17.26 ± 0.009	15.57 ± 0.012	14.43 ± 0.011
03 ^h 37 ^m 37 ^s 3	−36° 06′ 04″.7	18.24 ± 0.008	18.26 ± 0.030	18.29 ± 0.027	18.08 ± 0.025	16.91 ± 0.017	15.90 ± 0.017
03 ^h 38 ^m 47 ^s 5	−34° 28′ 20″.2	18.67 ± 0.016	18.27 ± 0.005	18.04 ± 0.002	17.65 ± 0.002	15.71 ± 0.030	14.21 ± 0.013
03 ^h 37 ^m 54 ^s 2	−33° 39′ 50″.5	19.09 ± 0.012	18.31 ± 0.004	17.72 ± 0.006	17.29 ± 0.014	15.87 ± 0.038	14.86 ± 0.010
03 ^h 38 ^m 00 ^s 5	−33° 34′ 17″.4	19.76 ± 0.016	18.33 ± 0.004	17.55 ± 0.003	17.19 ± 0.004	15.40 ± 0.089	14.34 ± 0.015
03 ^h 31 ^m 50 ^s 9	−34° 35′ 55″.6	19.35 ± 0.004	18.37 ± 0.005	17.85 ± 0.001	17.55 ± 0.021	15.87 ± 0.010	14.88 ± 0.010
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Table A4. The *UNKNOWN* catalogue (sample table, the full table is available online). This catalogue contains sources without radial velocity measurements in the literature. Columns from left- to right-hand represent RA (a), Declination (b), magnitude in *u*, *g*, *r*, *i*, *J*, and *Ks* (c–h). Quoted errors are statistical ones, systematic errors are not included.

RA (hms) (a)	Dec. (dms) (b)	<i>u</i> (AB mag) (c)	<i>g</i> (AB mag) (d)	<i>r</i> (AB mag) (e)	<i>i</i> (AB mag) (f)	<i>J</i> (Vega mag) (g)	<i>Ks</i> (Vega mag) (h)
03 ^h 31 ^m 51 ^s .2	−37° 02′ 56″.7	17.00 ± 0.021	16.07 ± 0.031	16.14 ± 0.052	16.17 ± 0.037	15.57 ± 0.009	15.65 ± 0.014
03 ^h 45 ^m 22 ^s .4	−35° 01′ 03″.4	17.20 ± 0.028	16.28 ± 0.040	16.18 ± 0.053	16.02 ± 0.076	15.11 ± 0.039	14.95 ± 0.009
03 ^h 43 ^m 25 ^s .3	−36° 30′ 07″.6	17.12 ± 0.025	16.31 ± 0.010	16.07 ± 0.015	15.95 ± 0.027	15.07 ± 0.021	14.89 ± 0.012
03 ^h 41 ^m 57 ^s .3	−34° 38′ 23″.9	17.35 ± 0.031	16.31 ± 0.026	16.24 ± 0.003	16.29 ± 0.010	15.75 ± 0.059	15.60 ± 0.018
03 ^h 39 ^m 56 ^s .3	−37° 08′ 13″.6	17.26 ± 0.016	16.34 ± 0.010	16.12 ± 0.005	16.11 ± 0.039	15.41 ± 0.009	15.23 ± 0.009
03 ^h 43 ^m 46 ^s .6	−33° 13′ 52″.5	17.16 ± 0.003	16.35 ± 0.006	16.16 ± 0.004	15.98 ± 0.002	15.13 ± 0.009	15.16 ± 0.016
03 ^h 42 ^m 04 ^s .4	−34° 10′ 55″.0	17.31 ± 0.029	16.35 ± 0.032	16.01 ± 0.036	16.01 ± 0.063	15.12 ± 0.009	14.79 ± 0.009
03 ^h 33 ^m 32 ^s .2	−33° 35′ 01″.8	16.96 ± 0.009	16.36 ± 0.005	16.11 ± 0.004	16.03 ± 0.002	15.38 ± 0.156	15.00 ± 0.014
03 ^h 41 ^m 27 ^s .2	−36° 15′ 13″.2	17.19 ± 0.015	16.36 ± 0.003	16.13 ± 0.015	15.92 ± 0.016	15.09 ± 0.006	14.94 ± 0.012
03 ^h 45 ^m 13 ^s .2	−35° 07′ 22″.7	17.34 ± 0.028	16.36 ± 0.040	16.22 ± 0.053	16.02 ± 0.076	14.98 ± 0.039	14.73 ± 0.007
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Table A5. Catalogue of the UCD/GC candidates (sample table, the full table is available online). Columns from left- to right-hand represent Candidate ID (a), RA (b), Declination (c), magnitude in *u*, *g*, *r*, *i*, *J*, and *Ks* (d–i). Quoted errors are statistical ones, systematic errors are not included.

ID (a)	RA (hms) (b)	DEC (dms) (c)	<i>u</i> (AB mag) (d)	<i>g</i> (AB mag) (e)	<i>r</i> (AB mag) (f)	<i>i</i> (AB mag) (g)	<i>J</i> (Vega mag) (h)	<i>Ks</i> (Vega mag) (i)
UCD-CAND-1	03 ^h 37 ^m 23 ^s .7	−33° 37′ 13″.2	19.79 ± 0.013	18.33 ± 0.004	17.66 ± 0.006	17.36 ± 0.014	16.06 ± 0.038	15.40 ± 0.014
UCD-CAND-2	03 ^h 33 ^m 24 ^s .8	−37° 21′ 40″.4	19.75 ± 0.008	18.58 ± 0.002	17.96 ± 0.004	17.65 ± 0.002	16.48 ± 0.016	16.05 ± 0.021
UCD-CAND-3	03 ^h 41 ^m 35 ^s .2	−33° 26′ 41″.8	19.79 ± 0.023	18.61 ± 0.021	18.02 ± 0.032	17.69 ± 0.021	16.50 ± 0.042	16.10 ± 0.035
UCD-CAND-4	03 ^h 43 ^m 54 ^s .4	−33° 54′ 10″.4	19.83 ± 0.059	18.64 ± 0.027	18.09 ± 0.081	17.87 ± 0.081	16.61 ± 0.017	15.84 ± 0.027
UCD-CAND-5	03 ^h 33 ^m 08 ^s .3	−37° 03′ 02″.7	19.81 ± 0.038	18.65 ± 0.038	18.12 ± 0.059	17.90 ± 0.012	16.76 ± 0.015	16.18 ± 0.023
UCD-CAND-6	03 ^h 46 ^m 52 ^s .4	−34° 15′ 24″.1	19.92 ± 0.006	18.65 ± 0.009	18.17 ± 0.012	17.85 ± 0.012	16.58 ± 0.025	15.82 ± 0.014
UCD-CAND-7	03 ^h 41 ^m 30 ^s .2	−34° 00′ 20″.4	20.23 ± 0.038	18.68 ± 0.031	18.09 ± 0.023	17.84 ± 0.050	16.45 ± 0.074	15.62 ± 0.024
UCD-CAND-8	03 ^h 48 ^m 16 ^s .3	−34° 22′ 01″.9	19.83 ± 0.007	18.68 ± 0.008	18.27 ± 0.008	17.95 ± 0.002	16.89 ± 0.018	16.27 ± 0.032
UCD-CAND-9	03 ^h 38 ^m 52 ^s .6	−35° 35′ 17″.7	19.95 ± 0.007	18.70 ± 0.010	18.25 ± 0.016	17.87 ± 0.029	16.99 ± 0.016	16.24 ± 0.014
UCD-CAND-10	03 ^h 32 ^m 51 ^s .7	−34° 03′ 01″.6	19.74 ± 0.016	18.72 ± 0.010	18.13 ± 0.007	17.93 ± 0.012	16.69 ± 0.044	16.22 ± 0.037
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GC-CAND-1	03 ^h 43 ^m 02 ^s .5	−34° 05′ 12″.8	21.99 ± 0.047	21.00 ± 0.038	20.50 ± 0.046	20.31 ± 0.064	19.10 ± 0.068	18.31 ± 0.229
GC-CAND-2	03 ^h 42 ^m 26 ^s .8	−33° 59′ 44″.1	22.44 ± 0.087	21.00 ± 0.044	20.91 ± 0.085	20.26 ± 0.081	19.13 ± 0.069	18.23 ± 0.216
GC-CAND-3	03 ^h 39 ^m 47 ^s .6	−36° 35′ 32″.7	22.03 ± 0.025	21.01 ± 0.003	20.61 ± 0.004	20.31 ± 0.024	19.39 ± 0.070	18.65 ± 0.217
GC-CAND-4	03 ^h 37 ^m 32 ^s .5	−36° 27′ 12″.9	22.26 ± 0.032	21.01 ± 0.006	20.50 ± 0.020	20.14 ± 0.025	19.01 ± 0.053	18.62 ± 0.211
GC-CAND-5	03 ^h 31 ^m 13 ^s .4	−36° 07′ 21″.6	22.31 ± 0.053	21.02 ± 0.059	20.26 ± 0.041	20.03 ± 0.054	18.93 ± 0.050	18.11 ± 0.164
GC-CAND-6	03 ^h 35 ^m 46 ^s .3	−35° 59′ 21″.4	22.27 ± 0.045	21.02 ± 0.047	20.42 ± 0.052	20.18 ± 0.031	18.94 ± 0.047	18.39 ± 0.203
GC-CAND-7	03 ^h 42 ^m 03 ^s .6	−34° 38′ 18″.3	22.66 ± 0.071	21.02 ± 0.032	20.15 ± 0.004	19.82 ± 0.009	17.89 ± 0.033	17.36 ± 0.094
GC-CAND-8	03 ^h 41 ^m 09 ^s .1	−36° 14′ 44″.1	22.05 ± 0.031	21.02 ± 0.007	20.56 ± 0.011	20.28 ± 0.018	19.12 ± 0.051	18.31 ± 0.159
GC-CAND-9	03 ^h 35 ^m 22 ^s .0	−35° 14′ 00″.0	22.36 ± 0.046	21.03 ± 0.009	20.46 ± 0.022	20.17 ± 0.015	18.74 ± 0.042	18.24 ± 0.136
GC-CAND-10	03 ^h 35 ^m 21 ^s .5	−35° 14′ 42″.1	22.23 ± 0.041	21.03 ± 0.009	20.53 ± 0.022	20.29 ± 0.016	18.95 ± 0.048	18.35 ± 0.151
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Table A6. Catalogue of the BEST UCD candidates. Columns from left- to right-hand represent candidate ID (a), RA (b), Declination (c), magnitude in u , g , r , i , J , and Ks (d–i), measured half-light (effective) radius (j). Quoted errors are statistical ones, systematic errors are not included.

ID	RA	Dec	u	g	r	i	J	Ks	r_h
–	(hms)	(dms)	(AB mag)	(AB mag)	(AB mag)	(AB mag)	Vega mag	(Vega mag)	(pc)
(a)	(b)	(c)	(d)	(e)	(f)	(g)	(h)	(i)	(j)
FDS-UCD-1	03 ^h 37 ^m 23 ^s .7	−33° 37′ 13 [″] .2	19.79 ± 0.013	18.33 ± 0.004	17.66 ± 0.006	17.36 ± 0.014	16.06 ± 0.038	15.40 ± 0.014	37.8
FDS-UCD-2	03 ^h 46 ^m 52 ^s .4	−34° 15′ 24 [″] .1	19.92 ± 0.006	18.65 ± 0.009	18.17 ± 0.012	17.85 ± 0.012	16.58 ± 0.025	15.82 ± 0.014	22.3
FDS-UCD-3	03 ^h 48 ^m 16 ^s .3	−34° 22′ 01 [″] .9	19.83 ± 0.007	18.68 ± 0.008	18.27 ± 0.008	17.95 ± 0.002	16.89 ± 0.018	16.27 ± 0.032	14.6
FDS-UCD-4	03 ^h 31 ^m 14 ^s .0	−33° 34′ 37 [″] .3	20.18 ± 0.017	18.98 ± 0.019	18.38 ± 0.004	18.14 ± 0.002	17.24 ± 0.184	16.37 ± 0.043	8.5
FDS-UCD-5	03 ^h 44 ^m 01 ^s .8	−33° 08 ^m 11 [″] .0	20.22 ± 0.007	19.09 ± 0.001	18.50 ± 0.001	18.16 ± 0.003	17.06 ± 0.018	16.49 ± 0.050	7.1
FDS-UCD-6	03 ^h 26 ^m 09 ^s .0	−36° 29′ 20 [″] .8	20.64 ± 0.012	19.32 ± 0.005	18.64 ± 0.002	18.32 ± 0.006	16.93 ± 0.015	16.12 ± 0.024	34.3
FDS-UCD-7	03 ^h 42 ^m 34 ^s .9	−34° 03′ 14 [″] .4	20.36 ± 0.030	19.35 ± 0.032	18.88 ± 0.036	18.63 ± 0.063	17.77 ± 0.030	16.84 ± 0.059	21.5
FDS-UCD-8	03 ^h 40 ^m 53 ^s .3	−33° 16′ 24 [″] .6	20.64 ± 0.083	19.47 ± 0.037	18.90 ± 0.053	18.59 ± 0.020	17.45 ± 0.022	16.88 ± 0.056	15.1
FDS-UCD-9	03 ^h 33 ^m 49 ^s .2	−37° 26′ 02 [″] .4	21.06 ± 0.012	19.77 ± 0.004	19.03 ± 0.004	18.71 ± 0.002	17.50 ± 0.023	16.85 ± 0.045	10.0
FDS-UCD-10	03 ^h 38 ^m 17 ^s .6	−33° 28′ 12 [″] .0	21.25 ± 0.025	19.78 ± 0.004	18.97 ± 0.003	18.74 ± 0.004	17.36 ± 0.092	16.44 ± 0.038	20.5
FDS-UCD-11	03 ^h 37 ^m 58 ^s .3	−33° 37′ 34 [″] .4	21.14 ± 0.021	19.95 ± 0.004	19.26 ± 0.006	18.96 ± 0.014	17.90 ± 0.046	17.18 ± 0.071	16.1
FDS-UCD-12	03 ^h 42 ^m 33 ^s .1	−34° 01′ 08 [″] .6	21.07 ± 0.032	19.95 ± 0.032	19.46 ± 0.036	19.13 ± 0.063	18.25 ± 0.039	17.48 ± 0.107	18.2
FDS-UCD-13	03 ^h 46 ^m 44 ^s .4	−34° 04′ 53 [″] .8	21.50 ± 0.020	20.00 ± 0.016	19.53 ± 0.077	19.13 ± 0.016	17.73 ± 0.125	17.22 ± 0.054	39.2
FDS-UCD-14	03 ^h 33 ^m 14 ^s .3	−35° 44′ 05 [″] .3	21.36 ± 0.020	20.01 ± 0.005	19.22 ± 0.020	19.03 ± 0.023	17.79 ± 0.025	17.17 ± 0.051	9.9
FDS-UCD-15	03 ^h 37 ^m 13 ^s .4	−33° 57′ 43 [″] .5	21.23 ± 0.022	20.05 ± 0.005	19.79 ± 0.022	19.25 ± 0.042	18.04 ± 0.033	17.29 ± 0.085	40.2
FDS-UCD-16	03 ^h 45 ^m 49 ^s .1	−33° 49′ 32 [″] .8	21.46 ± 0.059	20.06 ± 0.010	19.54 ± 0.062	19.17 ± 0.036	17.73 ± 0.031	17.14 ± 0.091	13.5
FDS-UCD-17	03 ^h 30 ^m 08 ^s .3	−36° 47′ 01 [″] .4	21.87 ± 0.034	20.18 ± 0.015	19.46 ± 0.012	19.07 ± 0.023	17.43 ± 0.021	16.61 ± 0.042	67.0
FDS-UCD-18	03 ^h 44 ^m 45 ^s .8	−36° 27′ 41 [″] .2	21.50 ± 0.039	20.20 ± 0.014	19.60 ± 0.024	19.33 ± 0.025	18.26 ± 0.036	17.53 ± 0.083	15.0
FDS-UCD-19	03 ^h 32 ^m 41 ^s .4	−33° 17′ 22 [″] .6	21.71 ± 0.037	20.30 ± 0.004	19.50 ± 0.005	19.24 ± 0.003	17.93 ± 0.025	17.24 ± 0.094	19.3
FDS-UCD-20	03 ^h 37 ^m 51 ^s .1	−33° 32′ 08 [″] .1	21.37 ± 0.026	20.30 ± 0.002	19.67 ± 0.003	19.44 ± 0.005	18.23 ± 0.087	17.59 ± 0.103	12.1
FDS-UCD-21	03 ^h 32 ^m 13 ^s .3	−35° 28′ 14 [″] .3	21.43 ± 0.021	20.32 ± 0.003	19.57 ± 0.006	19.52 ± 0.016	18.29 ± 0.033	17.67 ± 0.081	17.9
FDS-UCD-22	03 ^h 45 ^m 42 ^s .3	−35° 28′ 19 [″] .2	21.55 ± 0.018	20.32 ± 0.051	19.48 ± 0.099	19.11 ± 0.136	17.97 ± 0.028	17.19 ± 0.062	45.4
FDS-UCD-23	03 ^h 27 ^m 49 ^s .1	−34° 19′ 45 [″] .6	21.60 ± 0.025	20.34 ± 0.014	19.55 ± 0.003	19.29 ± 0.003	17.92 ± 0.034	17.40 ± 0.091	23.3
FDS-UCD-24	03 ^h 31 ^m 02 ^s .9	−34° 58′ 49 [″] .0	21.55 ± 0.026	20.38 ± 0.006	19.74 ± 0.029	19.51 ± 0.009	18.23 ± 0.037	17.55 ± 0.105	23.6
FDS-UCD-25	03 ^h 33 ^m 02 ^s .6	−33° 22′ 54 [″] .1	21.72 ± 0.037	20.39 ± 0.004	19.60 ± 0.005	19.32 ± 0.003	18.05 ± 0.027	17.33 ± 0.103	5.4
FDS-UCD-26	03 ^h 34 ^m 53 ^s .3	−33° 45′ 23 [″] .1	21.85 ± 0.037	20.48 ± 0.003	19.86 ± 0.011	19.59 ± 0.010	18.28 ± 0.035	17.45 ± 0.117	16.1
FDS-UCD-27	03 ^h 29 ^m 15 ^s .7	−34° 54′ 36 [″] .5	21.86 ± 0.039	20.54 ± 0.038	19.94 ± 0.027	19.65 ± 0.040	18.58 ± 0.046	17.83 ± 0.133	23.9
FDS-UCD-28	03 ^h 34 ^m 28 ^s .8	−37° 15′ 26 [″] .9	21.55 ± 0.017	20.56 ± 0.006	20.01 ± 0.003	19.80 ± 0.004	18.75 ± 0.048	18.05 ± 0.131	21.4
FDS-UCD-29	03 ^h 43 ^m 49 ^s .6	−36° 35′ 41 [″] .8	21.99 ± 0.027	20.58 ± 0.011	19.95 ± 0.017	19.60 ± 0.024	18.45 ± 0.043	17.34 ± 0.070	5.5
FDS-UCD-30	03 ^h 40 ^m 26 ^s .4	−34° 25′ 05 [″] .9	22.10 ± 0.035	20.63 ± 0.010	19.92 ± 0.005	19.63 ± 0.010	18.37 ± 0.043	17.71 ± 0.125	15.3
FDS-UCD-31	03 ^h 30 ^m 19 ^s .8	−35° 54′ 13 [″] .5	21.96 ± 0.053	20.67 ± 0.023	20.22 ± 0.070	19.96 ± 0.052	18.91 ± 0.049	18.14 ± 0.176	12.7
FDS-UCD-32	03 ^h 30 ^m 53 ^s .2	−37° 24′ 46 [″] .1	22.08 ± 0.027	20.73 ± 0.014	19.96 ± 0.040	19.67 ± 0.036	18.60 ± 0.038	17.60 ± 0.102	12.5
FDS-UCD-33	03 ^h 41 ^m 59 ^s .0	−36° 25′ 00 [″] .4	21.93 ± 0.031	20.78 ± 0.003	20.21 ± 0.013	19.94 ± 0.010	18.91 ± 0.045	18.16 ± 0.149	11.4
FDS-UCD-34	03 ^h 41 ^m 03 ^s .5	−37° 18′ 16 [″] .8	22.20 ± 0.027	20.80 ± 0.013	20.21 ± 0.008	19.85 ± 0.028	18.46 ± 0.037	17.65 ± 0.087	45.9
FDS-UCD-35	03 ^h 40 ^m 23 ^s .7	−36° 45′ 07 [″] .3	22.13 ± 0.033	20.82 ± 0.017	19.99 ± 0.009	19.68 ± 0.025	18.19 ± 0.032	17.34 ± 0.065	72.4
FDS-UCD-36	03 ^h 32 ^m 24 ^s .0	−35° 11′ 56 [″] .7	22.00 ± 0.035	20.84 ± 0.054	20.00 ± 0.012	19.84 ± 0.034	18.54 ± 0.087	17.85 ± 0.099	21.1
FDS-UCD-37	03 ^h 32 ^m 10 ^s .2	−35° 04′ 32 [″] .6	21.90 ± 0.032	20.84 ± 0.053	20.04 ± 0.010	19.82 ± 0.022	18.59 ± 0.100	17.76 ± 0.091	10.2
FDS-UCD-38	03 ^h 37 ^m 02 ^s .4	−33° 34′ 29 [″] .9	22.12 ± 0.039	20.85 ± 0.010	20.14 ± 0.004	19.78 ± 0.003	18.65 ± 0.085	17.74 ± 0.120	8.1
FDS-UCD-39	03 ^h 45 ^m 11 ^s .7	−35° 26′ 44 [″] .4	21.90 ± 0.025	20.87 ± 0.019	20.36 ± 0.013	19.99 ± 0.037	19.06 ± 0.053	18.37 ± 0.186	19.9
FDS-UCD-40	03 ^h 47 ^m 35 ^s .3	−35° 49′ 30 [″] .5	22.33 ± 0.042	20.91 ± 0.012	20.34 ± 0.024	19.90 ± 0.108	18.57 ± 0.038	17.52 ± 0.062	54.0
FDS-UCD-41	03 ^h 33 ^m 01 ^s .4	−36° 01′ 34 [″] .6	22.13 ± 0.037	20.91 ± 0.018	20.16 ± 0.022	19.92 ± 0.030	18.64 ± 0.041	17.97 ± 0.143	19.4
FDS-UCD-42	03 ^h 40 ^m 56 ^s .4	−33° 54′ 15 [″] .8	22.43 ± 0.058	20.91 ± 0.019	20.64 ± 0.052	19.89 ± 0.041	18.83 ± 0.053	17.86 ± 0.143	12.8
FDS-UCD-43	03 ^h 33 ^m 02 ^s .2	−35° 44′ 06 [″] .1	22.44 ± 0.049	20.99 ± 0.005	20.28 ± 0.013	19.99 ± 0.022	18.47 ± 0.037	17.72 ± 0.085	26.8
FDS-UCD-44	03 ^h 35 ^m 48 ^s .1	−35° 11′ 10 [″] .9	21.92 ± 0.034	20.99 ± 0.028	20.20 ± 0.028	20.03 ± 0.015	18.63 ± 0.054	18.18 ± 0.128	29.0

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