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# Excitation of vertical breathing motion in disc galaxies by tidally-induced spirals in fly-by interactions

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## ABSTRACT

It is now clear that the stars in the Solar neighbourhood display large-scale coherent vertical breathing motions. At the same time, Milky Way-like galaxies experience tidal interactions with satellites/companions during their evolution. While these tidal interactions can excite vertical oscillations, it is still not clear whether vertical breathing motions are excited *directly* by the tidal encounters or are driven by the tidally-induced spirals. We test whether excitation of breathing motions are directly linked to tidal interactions by constructing a set of  $N$ -body models (with mass ratio 5:1) of unbound single fly-by interactions with varying orbital configurations. We first reproduce the well-known result that such fly-by interactions can excite strong transient spirals (lasting for  $\sim 2.9$ – $4.2$  Gyr) in the outer disc of the host galaxy. The generation and strength of the spirals are shown to vary with the orbital parameters (the angle of interaction, and the orbital spin vector). Furthermore, we demonstrate that our fly-by models exhibit coherent breathing motions whose amplitude increases with height. The amplitudes of breathing motions show characteristic modulation along the azimuthal direction with compressing breathing motions coinciding with the peaks of the spirals and expanding breathing motions falling in the inter-arm regions – a signature of a spiral-driven breathing motion. These breathing motions in our models end when the strong tidally-induced spiral arms fade away. Thus, it is the tidally-induced spirals which drive the large-scale breathing motions in our fly-by models, and the dynamical role of the tidal interaction in this context is indirect.

**Key words:** methods: numerical – galaxies: evolution – galaxies: interaction – galaxies: kinematics and dynamics – galaxies: spiral – galaxies: structure.

## 1 INTRODUCTION

The second data release from the *Gaia* mission (hereafter *Gaia* DR2) has revealed the presence of large-scale bulk vertical motions ( $\sim 10$  km s<sup>-1</sup> in magnitude) and the associated bending and breathing motions for stars in the Solar vicinity and beyond (Gaia Collaboration et al. 2018). The presence of such breathing motions, i.e. stars on both sides of the Galactic mid-plane moving coherently towards or away from it, has also been reported in various past Galactic surveys, for example, the SEGUE (Sloan Extension for Galactic Understanding and Exploration) survey (Widrow et al. 2012), the LAMOST (Large Sky Area Multi-Object Fibre Spectroscopic Telescope) survey (Carlin et al. 2013), and the RAVE (Radial Velocity Experiment) data (Williams et al. 2013). The existence of such non-zero bulk vertical motions in the Milky Way raises questions about the plausible driving mechanism(s), since, in an axisymmetric potential, the bulk radial and vertical motions should be zero (e.g. Binney & Tremaine 2008).

Much of the understanding of the excitation of breathing motions in Milky Way-like galaxies have been gleaned from numerical simulations. Using semi-analytic models and test-particle simulations, Faure, Siebert & Famaey (2014) was the first to show that a strong spiral can drive large-scale vertical motions ( $|v_z| \sim 5$ – $20$  km s<sup>-1</sup>). The amplitude of such breathing motion increases at first with height from the mid-plane, and then starts to decrease after reaching its maximum value at a certain height. Also, using self-consistent  $N$ -body simulation, Debattista (2014) showed that a *vertically-extended* spiral feature can drive strong large-scale breathing motions with amplitude increasing monotonically from the mid-plane. The relative sense of these bulk motions, whether compressing or expanding changes across the corotation resonance (hereafter CR) of the spiral (Debattista 2014; Faure et al. 2014). Furthermore, using a *self-consistent*, high-resolution simulation with star formation, Ghosh, Debattista & Khachatryan (2022a) studied the age-dependence of such vertical breathing motions excited by spiral density waves. They showed that, at fixed height, the amplitude of such vertical breathing motion decreases with stellar age. They showed a similar age-variation in the breathing amplitude in the *Gaia* DR2, thereby

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supporting the scenario that the breathing motion of the Milky Way might well be driven by spiral density waves (Ghosh et al. 2022a). Instead, Monari, Famaey & Siebert (2015) showed that a stellar bar can also drive such breathing motion in disc galaxies. However, the resulting amplitudes of the breathing motions are small ( $|v_z| \sim 1 \text{ km s}^{-1}$ ) when compared to the spiral-driven breathing amplitudes.

In the Lambda cold dark matter ( $\Lambda$ CDM) paradigm of hierarchical structure formation, galaxies grow in mass and size via major mergers and/or multiple minor mergers, and cold gas accretion (White & Rees 1978; Fall & Efstathiou 1980). During the evolutionary phase, a galaxy also experiences multiple tidal interactions with satellites and/or passing-by companion galaxies. The frequency of such fly-by encounters increases at lower redshifts (e.g. see Sinha & Holley-Bockelmann 2015), and their cumulative dynamical impact on the morphology as well as on the dynamics of the host galaxies can be non-negligible (e.g. see An et al. 2019). Fly-by encounters can excite an  $m = 2$  bar mode (e.g. see Noguchi 1987; Łokas et al. 2016; Martínez-Valpuesta et al. 2017; Ghosh et al. 2021), off-set bars with a one-arm spiral (e.g. Pardy et al. 2016), and an  $m = 1$  lopsidedness in the stellar disc (Bournaud et al. 2005; Mapelli, Moore & Bland-Hawthorn 2008; Ghosh et al. 2022b). They can also trigger star formation (e.g. Duc, Cuillandre & Renaud 2018), and if the gas inflow towards the centre is exceedingly large, this can lead to a starburst (Mihos & Hernquist 1994) as well as triggering of AGN activity (Combes 2001). Furthermore, the role of fly-bys has been investigated in the context of forming warps (Kim et al. 2014; Senczuk et al. 2020), disc heating and disc thickening (Reshetnikov & Combes 1997; Kumar, Das & Kataria 2021) tidal bridges and streams (e.g. see Toomre & Toomre 1972; Duc & Renaud 2013) altering the galaxy spin (Choi & Yi 2017), and the evolution of classical and pseudo-bulges (Kumar et al. 2021). Our Galaxy has also experienced such a tidal interaction with the Sagittarius (Sgr) dwarf galaxy (e.g. see Majewski et al. 2003). Recent studies have indicated that such a tidal interaction could excite a ‘snail-shell’ structure (phase-space spiral), bending motions in the Solar neighbourhood (for details see e.g. Widrow et al. 2014; Antoja et al. 2018; Gaia Collaboration et al. 2018).

A tidal encounter with another galaxy can excite spiral features, as was first proposed in the seminal work of Holmberg (1941). Later, pioneering numerical work of Toomre & Toomre (1972) showed that a tidal interaction can excite tidal tails, bridges, and spiral features in the disc of the host galaxy for a wide variety of orbital configurations. Following that, numerical simulation has become an indispensable tool to study the dynamical effect of galaxy interactions. Several past studies attempted to understand the role of tidal encounters in the context of excitation of spirals as well as to understand the longevity and nature of the resulting spirals (e.g. see Sundelius et al. 1987; Donner & Thomasson 1994; Salo & Laurikainen 2000; Dobbs et al. 2010; Pettitt & Wadsley 2018). Furthermore, a recent study by Pettitt et al. (2017) investigated star formation and the properties of interstellar medium in tidally-induced spirals.

Tidal interactions can also induce vertical distortions and oscillations in the disc of the host galaxy (e.g. see Hunter & Toomre 1969; Araki 1985; Mathur 1990; Weinberg 1991; Vesperini & Weinberg 2000; Gómez et al. 2013; Widrow et al. 2014; D’Onghia et al. 2016). Widrow et al. (2014) proposed a dynamical scenario where a satellite galaxy, while plunging into the disc, can excite both bending and breathing motions. Interestingly, such tidal interactions also excited a strong spiral response within the disc in their model (see Fig. 10 there). Therefore, it is still unclear whether the tidal interactions are ‘directly’ responsible for driving breathing

motion, or the tidally-induced spirals are driving the breathing motions.

We aim to test this latter hypothesis in detail in this paper. We study a set of  $N$ -body models of galaxy fly-by interactions while varying the orbital parameters. We investigate the generation of the spiral features due to such fly-by encounters, and quantify the nature and longevity of such spirals in different fly-by models. We closely follow the generation of the vertical breathing motions and their subsequent evolution. In particular, we look for evidence that the generation and evolution of the vertical breathing motions is correlated with the temporal evolution of the tidally-induced spirals.

The rest of the paper is organized as follows. Section 2 provides the details of the simulation set-up and the fly-by models. Section 3 presents the results of the tidally-induced spirals, the density wave nature of spirals, and the temporal evolution of their strength. Section 4 measures the properties of the vertical breathing motions driven by these tidally-induced spirals. Section 5 discusses a few limitations of this work while Section 6 summarizes the main findings of this work.

## 2 SIMULATION SET-UP OF GALAXY FLY-BY MODELS

To motivate our study, we construct a set of  $N$ -body models of galaxy fly-bys where the host galaxy experiences an unbound interaction with a perturber galaxy. The mass ratio of the perturber and the host galaxy is set to 5:1, and is kept fixed for all the models considered here. A prototype of such a galaxy fly-by model is already presented in Kumar et al. (2021). Here, we construct a suite of fly-by models varying the orbital configuration (e.g. angle of interaction, orientation of the orbital spin vector). The details of the modelling and the simulation set-up is discussed in Kumar et al. (2021). For the sake of completeness, here we briefly mention the equilibrium model of the galaxies as well as the orbital configurations of the galaxy interaction.

### 2.1 Equilibrium models

The initial equilibrium model of each galaxy (host and the perturber) consists of a classical bulge, a stellar disc, and a dark matter (hereafter DM) halo. Each of the galactic components is treated as live, thereby allowing them to interact with each other. The DM halo is assumed to be spherically symmetric, and is modelled with a Hernquist density profile (Hernquist 1990) of the form

$$\rho_{\text{dm}}(r) = \frac{M_{\text{dm}}}{2\pi} \frac{a}{r(r+a)^3}, \quad (1)$$

where  $M_{\text{dm}}$  and  $a$  are the total mass and the scale radius of the DM halo, respectively. The scale radius of the Hernquist halo is related to the concentration parameter ‘ $c$ ’ of NFW DM halo (Navarro, Frenk & White 1996). For an NFW DM halo with mass  $M_{200} = M_{\text{dm}}$ , this relation is given by the following equation,

$$a = \frac{r_{200}}{c} \sqrt{2 \left[ \ln(1+c) - \frac{c}{(1+c)} \right]}, \quad (2)$$

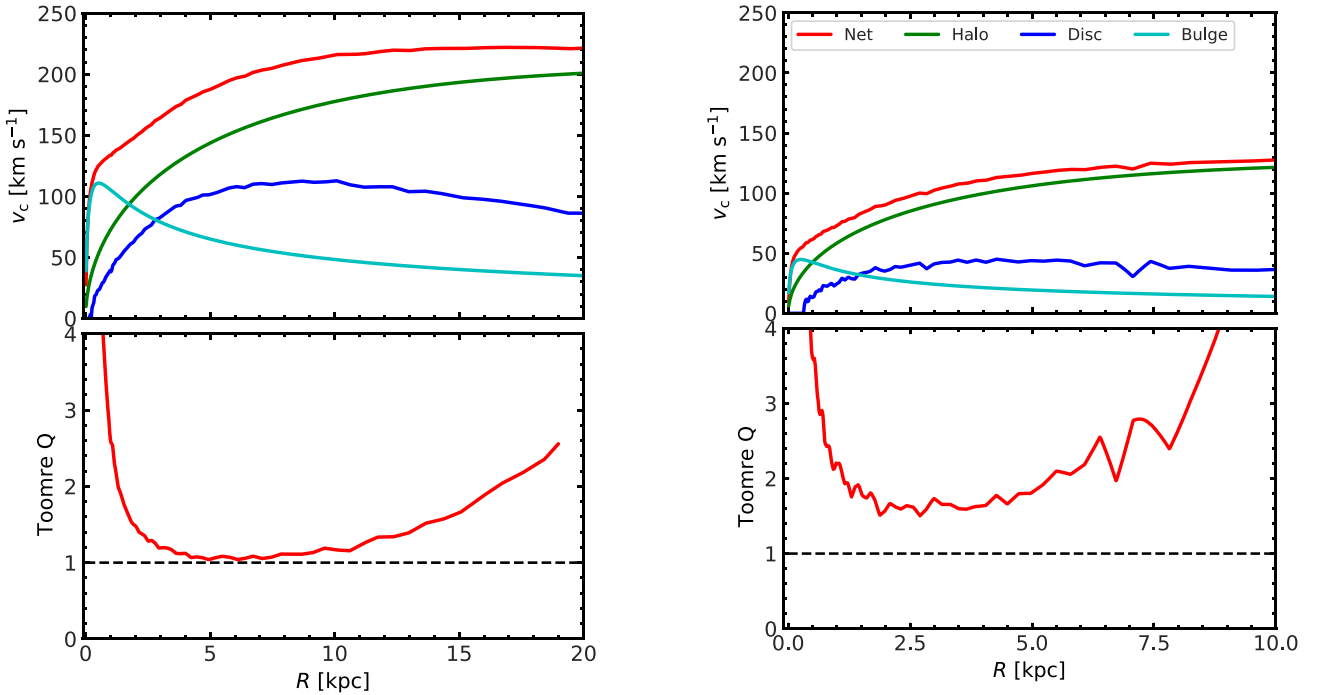
where  $r_{200}$  represents the radius of an NFW halo,<sup>1</sup> and the mass within this radius is defined as  $M_{200}$ . The classical bulge is also assumed

<sup>1</sup>It is defined as the radius from the centre of the halo inside which the mean density is 200 times the critical density of the Universe.

**Table 1.** Key galaxy parameters for the equilibrium models of the host and the perturber galaxies.

Galaxy	$M^{(1)}$ ( $\times 10^{12} M_{\odot}$ )	$\lambda^{(2)}$	$c^{(3)}$	$f_{\text{disc}}^{(4)}$	$f_{\text{bulge}}^{(5)}$	$j_d^{(6)}$	$R_d^{(7)}$ (kpc)	$z_0^{(8)}$ (kpc)	$N_{\text{halo}}^{(9)}$ ( $\times 10^6$ )	$N_{\text{disc}}^{(10)}$ ( $\times 10^6$ )	$N_{\text{bulge}}^{(11)}$ ( $\times 10^6$ )	$N_{\text{tot}}^{(12)}$ ( $\times 10^6$ )
Host	1.2	0.035	10	0.025	0.005	0.03	3.8	0.38	2.5	1.5	1	5
Perturber	0.24	0.035	11	0.01	0.002	0.01	1.95	0.195	0.5	0.3	0.2	1

*Note.* (1) total mass (in  $M_{\odot}$ ); (2) halo spin; (3) halo concentration; (4) disc mass fraction; (5) bulge mass fraction; (6) disc spin fraction; (7) disc scale length (in kpc); (8) disc scale height (in kpc); (9) total DM halo particles; (10) total disc particles; (11) total bulge particles; (12) total number of particles used.



**Figure 1.** Radial profiles of the circular velocity ( $v_c$ ) and the Toomre  $Q$  parameter are shown for equilibrium models of the host (left-hand panels) and the perturber galaxy (right-hand panels). In the top panels, the blue line denotes the contribution of the stellar disc while the DM halo contribution is shown by the green line. The bulge contribution is shown by the cyan solid line whereas the red line denotes the total/net circular velocity.

to be spherically symmetric and is modelled with another Hernquist density profile (Hernquist 1990) of the form

$$\rho_b(r) = \frac{M_b}{2\pi} \frac{b}{r(r+b)^3}, \quad (3)$$

where  $M_b$  and  $b$  represent the total bulge mass and the bulge scale radius, respectively. The initial radial surface density of the stellar disc follows an exponential fall-off and has a  $\text{sech}^2$  profile along the vertical direction, thereby having the form

$$\rho_d(R, z) = \frac{M_d}{4\pi z_0 R_d^2} \exp\left(-\frac{R}{R_d}\right) \text{sech}^2\left(\frac{z}{z_0}\right), \quad (4)$$

where  $M_d$  is the total mass,  $R_d$  is the exponential disc scale length, and  $z_0$  is the scale height. The corresponding values of the structural parameters used to model the host as well as the perturber galaxy are listed in Table 1.

The equilibrium models for the host as well as the perturber are generated using the publicly distributed code GALIC (Yurin & Springel 2014). This code uses elements of the Schwarzschild's method and the made-to-measure method to search for a stable solution of the collisionless Boltzmann Equation (CBE) for a set of collisionless stellar particles initialized by some predefined analytic

density distribution functions (for details, see Yurin & Springel 2014). A total of  $5 \times 10^6$  particles are used to model the host galaxy whereas a total of  $1 \times 10^6$  particles are used to model the perturber galaxy. The number of particles used to model each of the galaxy components of the host and the perturber galaxy are also listed in Table 1. The stellar particles have gravitational softening  $\epsilon = 20$  pc while the DM halo particles have  $\epsilon = 30$  pc. Fig. 1 shows the corresponding radial profiles of the circular velocity ( $v_c$ ) and the Toomre  $Q$  parameter for the host and the perturber galaxy at  $t = 0$ .

## 2.2 Set-up of galaxy fly-by scenario

To simulate the unbound galaxy fly-by scenario, we place our galaxy models on a hyperbolic orbit with eccentricity,  $e = 1.1$  so that the orbit of the perturber galaxy remains unbound throughout the interaction. We avoid choosing a parabolic orbit as the dynamical friction of the host galaxy decays the orbit of the perturber galaxy and puts the perturber galaxy on a bound elliptical orbit. Our choice of hyperbolic orbit avoids a bound fly-by interaction. We place the galaxies at an initial separation of 255 kpc before the start of the simulation. For different models with different orbital configurations,

**Table 2.** Key orbital parameters for the galaxy fly-by models.

Model <sup>(a)</sup>	$r_p^{(b)}$ (kpc)	$i^{(c)}$ (deg.)	$t_p^{(d)}$ (Gyr)	$t_{\text{enc}}^{(e)}$ (Gyr)	$T_p^{(f)}$
<i>RP40i00pro</i>	53.09	0	0.88	0.1298	-4.1348
<i>RP40i30pro</i>	52.34	30	0.85	0.1272	-4.1162
<i>RP40i60pro</i>	52.14	60	0.85	0.1265	-4.1113
<i>RP40i90pro</i>	52.23	90	0.85	0.1267	-4.1133
<i>RP40i00ret</i>	53.07	0	0.85	0.1298	-4.1342

Note. (a) Galaxy fly-by model; (b) pericentre distance (in kpc); (c) orbital angle of interaction (in degree); (d) time of pericentre passage (in Gyr); (e) encounter time (in Gyr); (f) tidal parameter.

we vary the distance of their closest approach (the pericentre distance) assuming the two-body Keplerian orbit. For further details of the orbital configuration and the geometry of the unbound fly-by scenario, the reader is referred to Kumar et al. (2021). A total of five such galaxy fly-by models are used for this study. The angle of interaction, and the pericentre distance for these models are listed in Table 2.

All the simulations are run using the publicly available code GADGET-2 (Springel, Yoshida & White 2001; Springel 2005a, b) for a total time of 6 Gyr with a tolerance parameter  $\theta_{\text{tol}} = 0.5$  and an integration time-step 0.4 Myr. The maximum error in the angular momentum of the system is well within 0.15 per cent throughout the evolution for all models considered here.

Following Di Matteo et al. (2007), we also calculate the encounter time ( $t_{\text{enc}}$ ), and tidal parameter ( $T_p$ ) for the host galaxy by using

$$t_{\text{enc}} = \frac{r_p}{v_p}, T_p = \log_{10} \left[ \frac{m_{\text{per}}}{M_{\text{host}}} \left( \frac{R_d}{r_p} \right)^3 \right], \quad (5)$$

where  $r_p$  is pericentre distance and  $v_p$  is the corresponding relative velocity.  $M_{\text{host}}$  and  $m_{\text{per}}$  are the masses of the host and the perturber galaxy, respectively. The corresponding values of  $t_{\text{enc}}$ , and  $T_p$  for all the models consider here are listed in Table 2. We note that the quantities mentioned here will be in units of the dimensionless parameter ‘ $h$ ’, defined via the Hubble constant ( $H_0 = 100 \text{ h km s}^{-1} \text{ Mpc}^{-1}$ ) and can be scaled to observed values.

Each model is referred as a unique string given by ‘[PERICENTRE DISTANCE][ANGLE OF INTERACTION][ORBITAL SPIN]’ where [PERICENTRE DISTANCE] denotes the pericenter distance, obtained using the standard two-body formalism. [ANGLE OF INTERACTION] denotes the angle at which the perturber encounters with the host galaxy while [ORBITAL SPIN] denotes the orbital spin vector (‘pro’ for prograde and ‘ret’ for retrograde orbits). We follow this scheme of nomenclature throughout the paper. As an example, RP40i30pro denotes a fly-by model where the perturber galaxy interacts with the host galaxy at an angle of  $30^\circ$  in a prograde orbit, and the calculated pericentre distance between these two galaxies is 40 kpc, obtained by using the standard two-body Keplerian orbit.

### 3 QUANTIFICATION OF TIDALLY-INDUCED SPIRALS IN GALAXY FLY-BYS

Here, we investigate the excitation of spiral structure in the host galaxy as a consequence of the tidal interaction with the perturber galaxy. We first study the model *RP40i00pro*, where the host galaxy experiences an in-plane unbound fly-by interaction with the perturber galaxy (with mass 1/5th that of the host galaxy). During the pericentre passage, the perturber exerts a strong tidal pull on the host galaxy. Fig. 2 (top panels) shows the face-on density distribution of the stars

of the host galaxy at four different times after the tidal interaction occurs. A visual inspection reveals that after the interaction happens, a strong spiral feature is excited in the outer parts of the host galaxy (e.g. at  $t = 2$  Gyr). The spiral features are also seen at later times. However, by the end of the simulation ( $t = 6$  Gyr), there is no discernible, strong spiral features left in the host galaxy. To study this trend further, we calculate the residual surface density ( $\tilde{\Sigma}(R, \phi)$ ) in the  $(R, \phi)$ -plane using

$$\tilde{\Sigma}(R, \phi) = \frac{\Sigma(R, \phi) - \Sigma_{\text{avg}}(R)}{\Sigma_{\text{avg}}(R)}, \quad (6)$$

where  $\Sigma_{\text{avg}}(R)$  is the azimuthally-averaged surface density of the disc at radius  $R$ . This is shown in Fig. 2 (bottom panels). As seen clearly, after the interaction happens, a strong, coherent spiral feature, denoted by the presence of a periodic over- and under-density, is excited in the outer region ( $R \geq 10$  kpc) of the host galaxy. At the end of the simulation ( $t = 6$  Gyr), the corresponding residual density distribution does not exhibit any coherent spiral structure in the disc of the host galaxy.

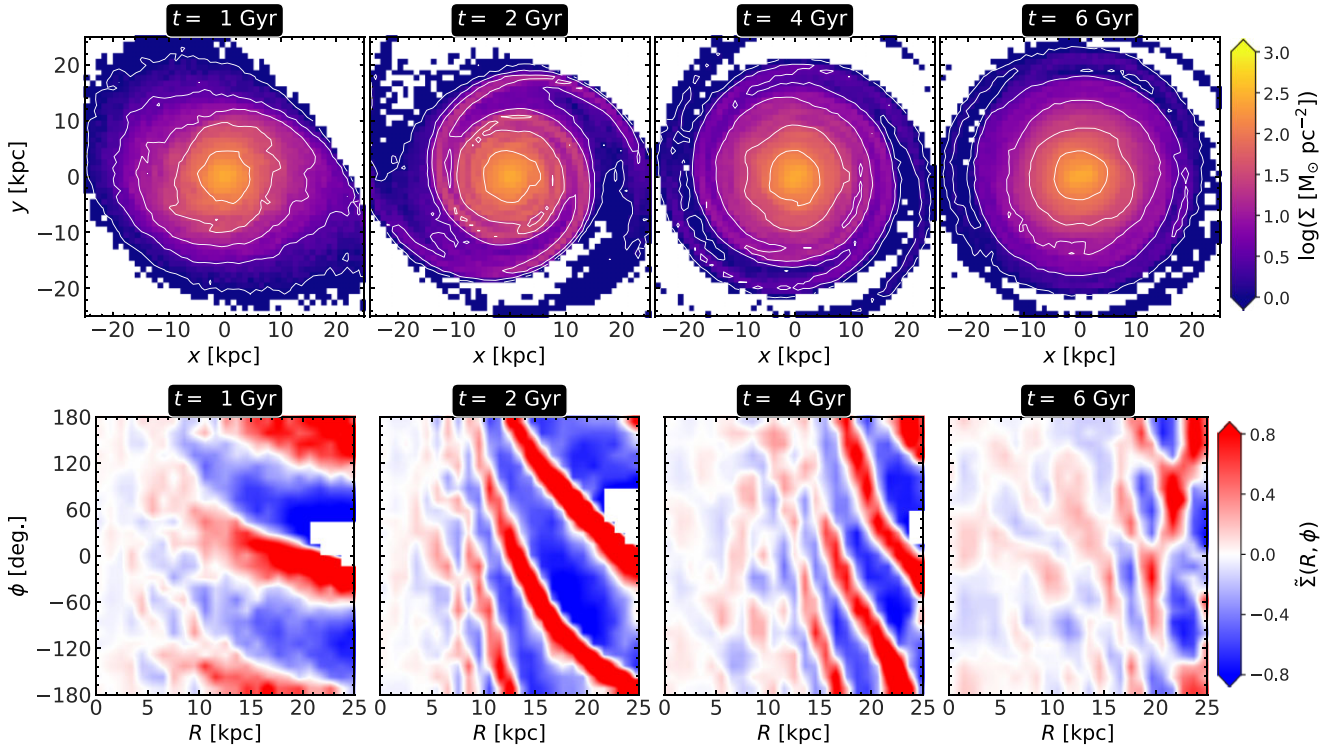
#### 3.1 Strength and temporal evolution of spirals

In the previous section, we have shown that a tidal interaction with a perturber galaxy excites a prominent spiral feature in the outer disc region of the host galaxy for the model *RP40i00pro*. Next, we quantify the strength of the spiral and follow its temporal evolution. For this, we first calculate the radial variation of the Fourier moment of the surface density of the stellar particles of the host galaxy using

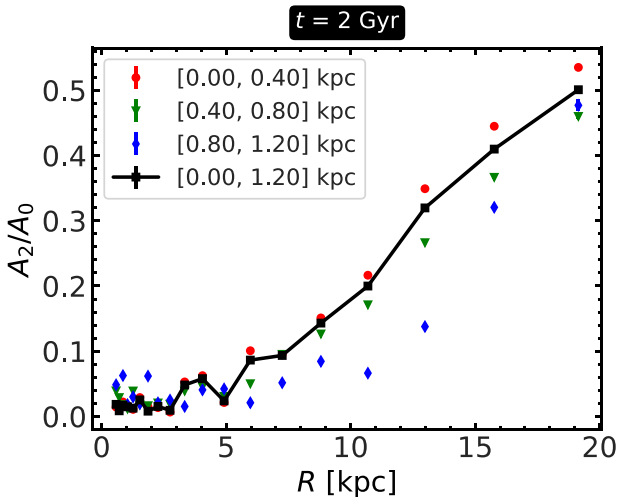
$$A_m/A_0(R) = \left| \frac{\sum_j m_j e^{im\phi_j}}{\sum_j m_j} \right|, \quad (7)$$

where  $A_m$  is the coefficient of the  $m$ th Fourier moment of the density distribution,  $m_j$  is the mass of the  $j$ th particle (e.g. see Kumar et al. 2021; Kumar, Das & Kataria 2022). Fig. 3 shows the corresponding radial variation of the  $m = 2$  Fourier moment at  $t = 2$  Gyr for the model *RP40i00pro*. In the outer disc regions, there are less particles when compared to the inner disc regions. Therefore, using a *linearly-spaced* radial binning when calculating the Fourier coefficient (using equation 7) introduces noise in the calculation for the outer regions. To avoid that, we employ a logarithmic binning in the radial direction. As seen clearly, in the outer parts ( $R \geq 10$  kpc), the values of the coefficient  $A_2/A_0$  are non-zero, indicating the presence of a strong spiral structure. Are these tidally-induced spirals mostly confined to the disc mid-plane or are they vertically-extended? To investigate this further, we calculate the radial variation of the same Fourier coefficient  $A_2/A_0$ , but for stars in different vertical layers of thickness 400 pc. The resulting radial variations are also shown in Fig. 3. As seen clearly from Fig. 3, the Fourier coefficient  $A_2/A_0$  shows non-zero values even for stars at the largest heights from the mid-plane ( $|z| = [0.8, 1.2]$  kpc). Also, at a certain radius  $R$  within the extent of the spirals, the values of the Fourier coefficient  $A_2/A_0$  decreases monotonically as one moves farther away from the mid-plane. We checked this variation of the Fourier coefficient  $A_2/A_0$  with height at other time-steps as well, and found that this trend remains generic whenever the tidally-induced spirals are strong in the disc of the host galaxy. This demonstrates that the tidally-induced spirals in our model *RP40i00pro* is *vertically-extended*, similar to what was reported in Debattista (2014) and Ghosh et al. (2022a).

To further investigate the spatio-temporal evolution of the tidally-induced spirals in model *RP40i00pro*, we calculate the Fourier coefficient  $A_2/A_0(R)$  at different radial locations for the whole

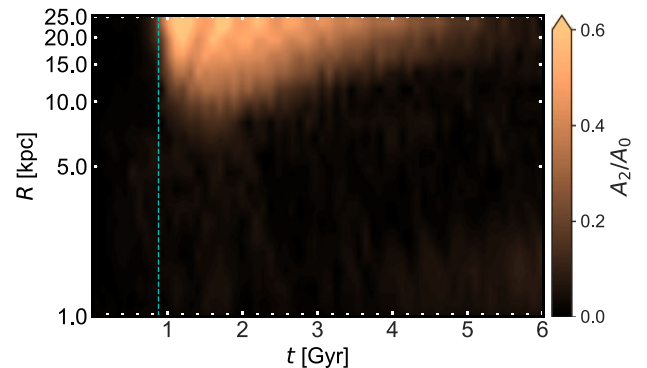


**Figure 2.** Top: the density distribution of the stars of the host galaxy in the face-on projection at four different times after the pericentre passage for the model *RP40i00pro*. White solid lines show the contours of constant surface density. A strong spiral feature is excited in the disc of the host galaxy after it experiences a tidal encounter with the perturber galaxy. Bottom: the corresponding distribution of the residual surface density ( $\tilde{\Sigma}(R, \phi)$ ) calculated using equation (6) at the same four different times. The sense of rotation is towards increasing  $\phi$ .



**Figure 3.** Radial variation of the Fourier coefficient of the  $m = 2$ . Fourier component (normalized by the  $m = 0$  component) is shown at  $t = 2$  Gyr for the model *RP40i00pro* (see the black solid line). The same quantity is then measured for stars in different vertical layers (as indicated in the legend).

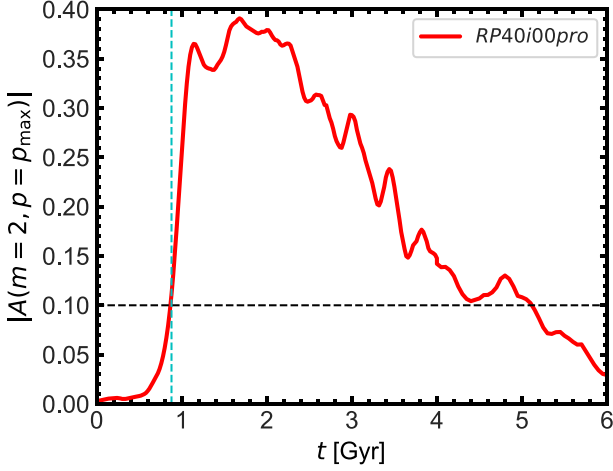
simulation run-time (a total of 6 Gyr). This is shown in Fig. 4. The tidal interaction with the perturber excites a strong spiral feature in the disc of the host galaxy after  $t \sim 1.1$  Gyr or so. These spirals remain mostly in the outer regions of the host’s disc ( $R \geq 10$  kpc), as the values of  $A_2/A_0$  in the inner part ( $R \leq 10$  kpc) are almost zero



**Figure 4.** The evolution of the coefficient of the  $m = 2$  Fourier moment ( $A_2/A_0$ ) in the  $R - t$  space for the model *RP40i00pro*. A logarithmic binning is employed along the radial direction, for details see text. The colour bar shows the values of  $A_2/A_0$ . The vertical dashed line (in cyan) denotes the time of the pericentre passage of the perturber.

(see Fig. 4). After  $t \sim 3$  Gyr or so, the spiral starts to decay, as shown by the decreasing values of the  $A_2/A_0$ . By the end of the simulation run, the values of the  $A_2/A_0$  become almost zero, implying that the tidally-induced spirals have wound up almost completely.

Next, we quantify the temporal evolution of the strength of the tidally-induced spirals in our fly-by model *RP40i00pro*. Following Sellwood & Carlberg (1984a), Sellwood & Athanassoula (1986),



**Figure 5.** Temporal evolution of the spiral strength ( $|A(m=2, p=p_{\max})|$ ), calculated using equation (8) for the model *RP40i00pro*. The tidally-induced spirals decay by the end of the simulation run. The horizontal black dotted line denotes  $|A(m=2, p=p_{\max})| = 0.1$ , used as an operational definition for the onset of the spirals, see text for details. The vertical dashed line (in cyan) denotes the time of the pericentre passage of the perturber.

Puerari et al. (2000), we define

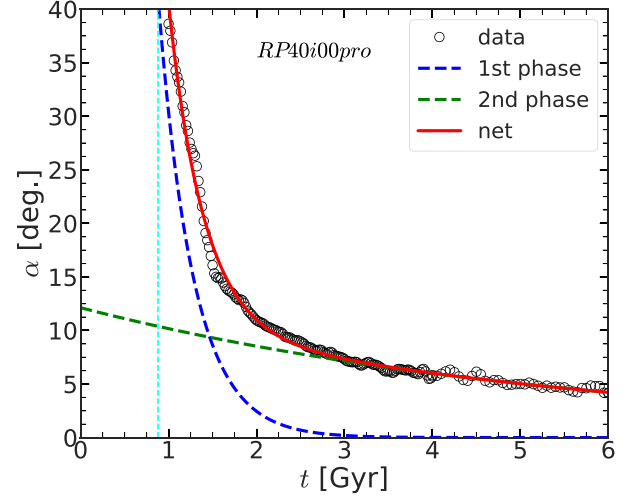
$$A(m, p) = \frac{\sum_j^N m_j \exp[i(m\phi_j + p \ln R_j)]}{\sum_j^N m_j}, \quad (8)$$

where  $|A(m, p)|$  is the amplitude,  $m_j$  is the mass of  $j^{\text{th}}$  star,  $m$  is the spiral arm multiplicity,  $(R_j, \phi_j)$  are the polar coordinates of the  $j^{\text{th}}$  star in the plane of the disc, and  $N$  is the total number of stellar particles in the annulus  $R_{\min} \leq R \leq R_{\max}$  within which the spiral feature exists and/or is most prominent. Here, we take  $R_{\min} = 3R_d$ , and  $R_{\max} = 6R_d$ , where  $R_d = 3.8$  kpc. For this annular region, we estimate  $A(m, p)$  as a function of  $p$  for  $p \in [-50, 50]$  with a fixed step of  $dp = 0.25$  (as suggested by Puerari et al. 2000) for different values of  $m$ . Then, we evaluate the parameter  $p_{\max}$  which corresponds to the maximum value of  $|A(m, p)|$ . We find that the amplitude  $|A(m, p=p_{\max})|$  shows a maximum value for  $m = 2$ , indicating that the  $m = 2$  spiral is the strongest. Therefore, at a certain time  $t$ , we define the amplitude  $|A(m=2, p=p_{\max})|$  as the strength of spirals (Oh, Kim & Lee 2015a; Semiczuk, Łokas & del Pino 2017; Kumar et al. 2021). The resulting temporal evolution of the strength of spirals for the model *RP40i00pro* is shown in Fig. 5. As seen from Fig. 5, the tidally-induced spirals grow for some time after the interaction happens, then remain stable for about 1 Gyr before weakening from around  $t = 3$  Gyr. By the end of the simulation, the spirals' strength becomes almost zero. For quantifying the longevity of the spirals, we define  $|A(m=2, p=p_{\max})| = 0.1$  as the onset of the strong spiral perturbation.<sup>2</sup> Therefore, the spirals persist for a time of  $\sim 4.2$  Gyr (after their formation) for the model *RP40i00pro* (also see Fig. 5).

### 3.2 Winding of the tidally-induced spirals

Lastly, we investigate the winding of the tidally-induced spirals in our fly-by model *RP40i00pro*. As shown in past  $N$ -body simulations of galactic discs (e.g. see Oh et al. 2008; Struck, Dobbs & Hwang 2011;

<sup>2</sup>The threshold value of 0.1 is used purely as an operational definition for the onset of the spirals.



**Figure 6.** Temporal evolution of the pitch angle ( $\alpha$ ) of the tidally-induced spirals is shown for the model *RP40i00pro*. A double-exponential profile (see equation 9) is fitted to model the temporal evolution. The blue dashed line denotes the initial rapid winding phase whereas the green dashed line denotes the subsequent slow winding phase, for details see text. The red solid line denotes the best-fit double-exponential profile. The vertical dashed line (in cyan) denotes the time of the pericentre passage of the perturber.

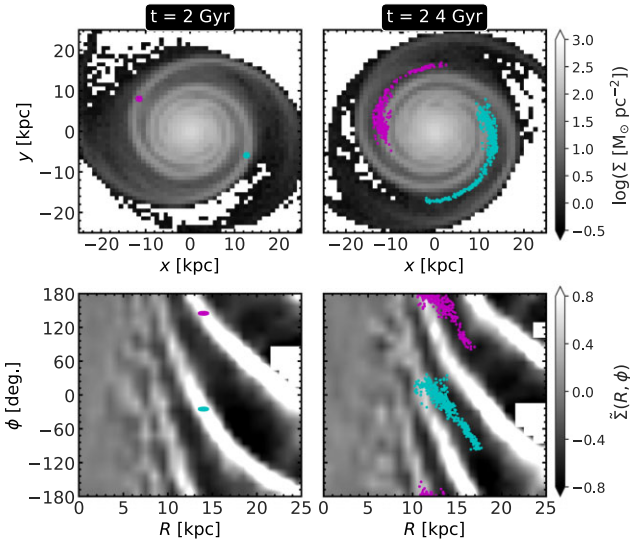
Kumar et al. 2021), a spiral arm can wind up with time. Following Oh et al. (2015a) and Semiczuk et al. (2017) at a certain time  $t$ , we define the pitch angle,  $\alpha$ , as  $\alpha = \tan^{-1}(m/p_{\max})$  using equation (8). The resulting temporal evolution of the pitch angle for the model *RP40i00pro* is shown in Fig. 6. As revealed in Fig. 6, the temporal evolution of the pitch angle displays two distinct phases, namely, the initial rapid winding phase where the pitch angle decreases sharply, and the subsequent slow winding phase where the pitch angle decreases less drastically. To model the temporal evolution of the pitch angle, we fit a double-exponential profile having the form

$$\alpha = \alpha_1 \exp[-\lambda_1(t - t_0)] + \alpha_2 \exp[-\lambda_2(t - t_0)], \quad (9)$$

where  $\alpha_1, \alpha_2, \lambda_1$ , and  $\lambda_2$  are free parameters. Here,  $t_0 = 1$  Gyr, and denotes the time of the spirals' formation. The fitting is performed via the SCIPY package CURVEFIT which uses the Levenberg–Marquardt algorithm. The resulting best-fit double exponential profile is shown in Fig. 6. We define the winding time-scale,  $\tau_{\text{wind}}$ , as  $\tau_{\text{wind}} = \alpha/|\dot{\alpha}|$  where  $|\dot{\alpha}|$  denotes the (absolute) rate of change of the pitch angle with time. For the initial rapid winding phase,  $\tau_{\text{wind}} \simeq 1/\lambda_1$  whereas for the subsequent slow winding phase  $\tau_{\text{wind}} \simeq 1/\lambda_2$  (see equation 9). We find the best-fitting values as  $\lambda_1 = 2.51 \pm 0.05 \text{ Gyr}^{-1}$  and  $\lambda_2 = 0.18 \pm 0.01 \text{ Gyr}^{-1}$ , which translate to a winding time-scale  $\tau_{\text{wind}} = 0.4$  Gyr for the initial rapid winding phase, and a winding time-scale  $\tau_{\text{wind}} = 5.7$  Gyr for the subsequent slow winding phase for the model *RP40i00pro*.

### 3.3 Nature of tidally-induced spirals

So far, we have demonstrated that a fly-by interaction with a perturber induces a strong spiral feature in the outer disc of the host galaxy. However, the question remains whether these spirals are density waves or material arm in nature. For a comprehensive review on the nature of the spirals in disc galaxies, the reader is referred to Binney & Tremaine (2008), and Dobbs & Baba (2014).

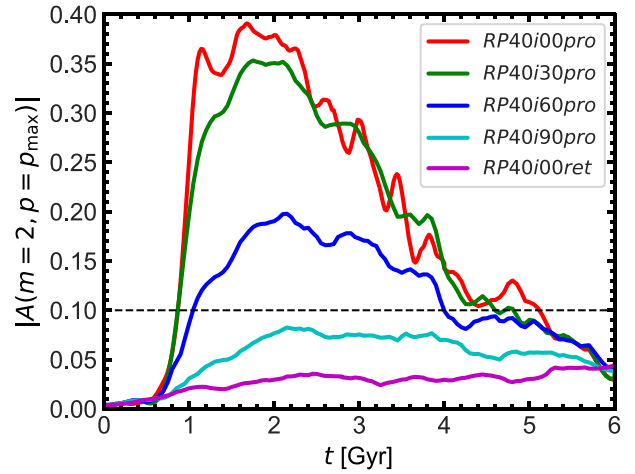


**Figure 7.** *Nature of spirals:* top and bottom panels of the first column show the surface density and residual density maps of *RP40i00pro* model at  $t = 2$  Gyr over-plotted with two patches (in magenta and cyan dots) of stellar particles in fully developed spiral arms. Right column is the analogous to the left column, but for  $t = 2.4$  Gyr. Both patches of stars, initially associated along the spirals have subsequently sheared out by the differential rotation, and left the spirals.

For the set of the  $N$ -body models we are using here, we do not have the age information of the stellar particles. This, in turn, restricts us from dividing stellar particles into different age-bins to trace the existence of spirals in different stellar population with different ages, as previously done in Ghosh et al. (2022a). Therefore, following Grand, Kawata & Cropper (2012), D’Onghia, Vogelsberger & Hernquist (2013), we test the nature of the tidally-induced spirals by following the stars which are located on spiral arms at a certain time. We chose a time, say  $t = 2$  Gyr when the spirals are fully-developed for the model *RP40i00pro*, and select two small patches (shown in magenta and cyan) of stars along the arms. This is shown in Fig. 7 (top left-hand panel). Now, if the spirals are of material arm in nature, then the stars would not leave the spiral arm at a subsequent time. To check that, we follow the selected stars at a later time,  $t = 2.4$  Gyr (see top right-hand panel of Fig. 7). As seen clearly from Fig. 7, the stars, initially concentrated in small patches along the spiral arm have sheared out due to the underlying differential rotation. This shearing out of the stars is more prominent in the distribution shown in the  $(R, \phi)$ -plane (see bottom panels of Fig. 7). Interestingly, the stars, initially contained in the patches have left the spirals at a subsequent time, and the pitch angle of the selected stars is different from that of the spirals at  $t = 2.4$  Gyr. Furthermore, we calculate the pattern speed of the spirals at different radial locations in the disc. We find that the pattern speeds of the spirals are very different from the values of the circular frequency ( $\Omega$ ) calculated at different radial locations. For brevity, these are not shown here. This shows that the spirals present in our fly-by model *RP40i00pro* are density waves in nature.

### 3.4 Dependence on orbital parameters

Here, we explore the generation of the spirals in the host galaxy due to a tidal interaction with the perturber galaxy for different angles of interaction of the tidal encounter as well as for different orbital spin vectors (prograde or retrograde). Intuitively, the response



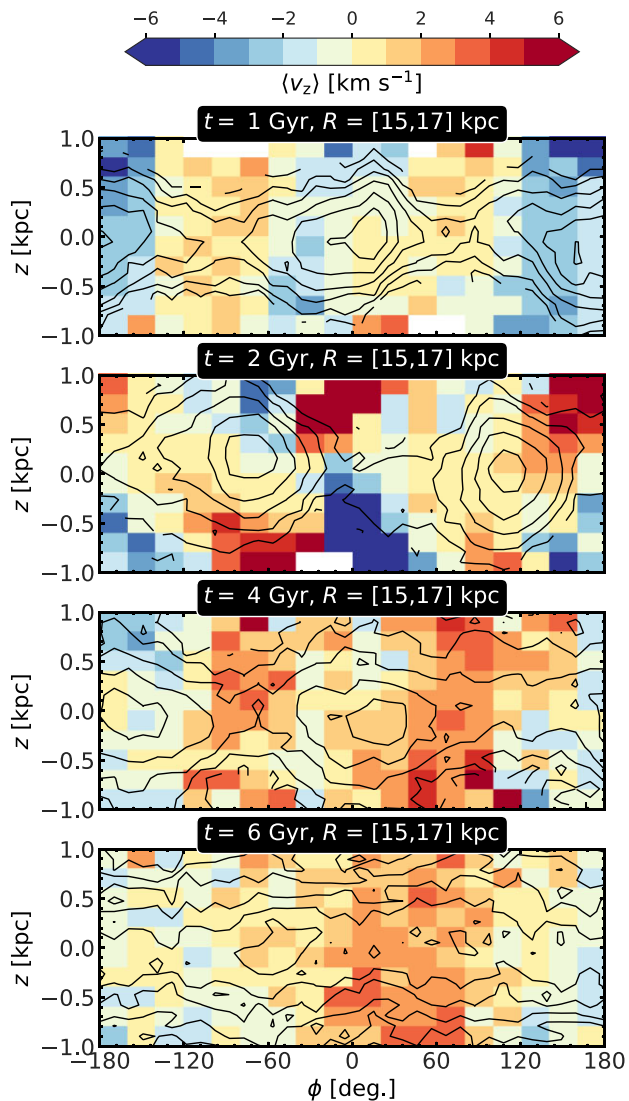
**Figure 8.** Temporal evolution of the spiral strength ( $|A(m = 2, p = p_{\max})|$ ), calculated using equation (8) is shown for models with different angle of interaction, and orbital spin vector. For comparison, we kept the model *RP40i00pro* here (see red solid line). For models with prograde configuration, the strength of spirals decreases monotonically from the co-planar ( $i = 0^\circ$ ) to polar ( $i = 90^\circ$ ) orbital configuration. The horizontal black dotted line denotes  $|A(m = 2, p = p_{\max})| = 0.1$ , used as an operational definition for the onset of the spirals.

of the host galaxy will be different for polar ( $i = 90^\circ$ ) and co-planar ( $i = 0^\circ$ ) orbital configurations. First we consider the other prograde models with angles of interaction  $30^\circ$  and  $60^\circ$  (for details see Section 2.2). Both the models exhibit a similar trend of excitation of tidally-induced spirals in the disc of the host galaxy as in model *RP40i00pro*. Shortly after the tidal encounter happens, the disc of the host develops a spiral feature which grows for a certain time, and then starts decaying. We calculate the strength of the spirals using equation (8) at different times for these models. The resulting temporal variations of the strength of the spirals for these three models are shown in Fig. 8. The maximum strength of the tidally-induced spirals decreases monotonically with larger angle of interaction. Next, we use  $|A(m = 2, p = p_{\max})| = 0.1$  for defining the onset of the strong spiral perturbation. We find that the spirals persist for a time-scale of  $\sim 2.9$ – $4.2$  Gyr, depending on the angle of interaction (also see Fig. 8). For the polar ( $i = 90^\circ$ ) configuration, the spirals are very weak as the values of  $|A(m = 2, p = p_{\max})|$  remain below 0.1 (see Fig. 8).

Next, we study the strength of the spirals in a galaxy fly-by model where the perturber interacts with the host galaxy in a retrograde orientation (with  $i = 0^\circ$ ) (for details see Section 2.2). A visual inspection of the face-on distribution of the stars (in the host galaxy) does not reveal any prominent spirals. We again calculate the strength of the spirals using equation (8) for the model *RP40i00ret*, and show this also in Fig. 8. The value of the amplitude  $|A(m = 2, p = p_{\max})|$  remains close to zero throughout the evolution of the model *RP40i00ret*, indicating that no prominent spirals are triggered/excited by the tidal encounter with the perturber galaxy for this model.

Furthermore, we have checked that the tidally-induced spirals display a similar winding, namely, an initial rapid winding phase, followed by a slow winding phase, as seen for the model *RP40i00pro* (see Fig. 6). Also, we have checked the nature of the resulting spirals for these other models, using the same technique employed in Section 3.3. We find that for the models showing prominent spirals, the resulting spirals show a density wave nature, similar to the model





**Figure 9.** Distribution of the mean vertical velocity,  $\langle v_z \rangle$  in the  $(\phi, z)$ -plane for stars in the radial annulus  $15 \leq R/\text{kpc} \leq 17$  are shown at four different times for the model *RP40i00pro*. The solid black lines denote the contours of constant density. The presence of large-scale non-zero vertical velocities for stellar particles are seen at all four time-steps for details see text.

*RP40i00pro*. For the sake of brevity, we do not show these results here.

#### 4 BREATHING MOTIONS EXCITED BY TIDALLY-INDUCED SPIRALS

In the previous section, we have demonstrated that a tidal interaction with a perturber galaxy can excite prominent spirals in the disc of the host galaxy. Here, we investigate the dynamical impact of these tidally-induced spirals on the bulk vertical motions of the host galaxy. We first choose the fly-by model *RP40i00pro* which harbours a strong spiral after the interaction. We choose the radial extent  $15 \leq R/\text{kpc} \leq 17$  where the spirals are prominent at later times in this radial annulus, and we calculate the mean vertical velocity ( $\langle v_z \rangle$ ) in the  $(\phi, z)$ -plane at four different times, namely, at  $t = 1, 2, 4,$  and  $6$  Gyr (same as in Fig. 2). This is shown in Fig. 9. During our chosen time interval, spirals' strength varies from strong

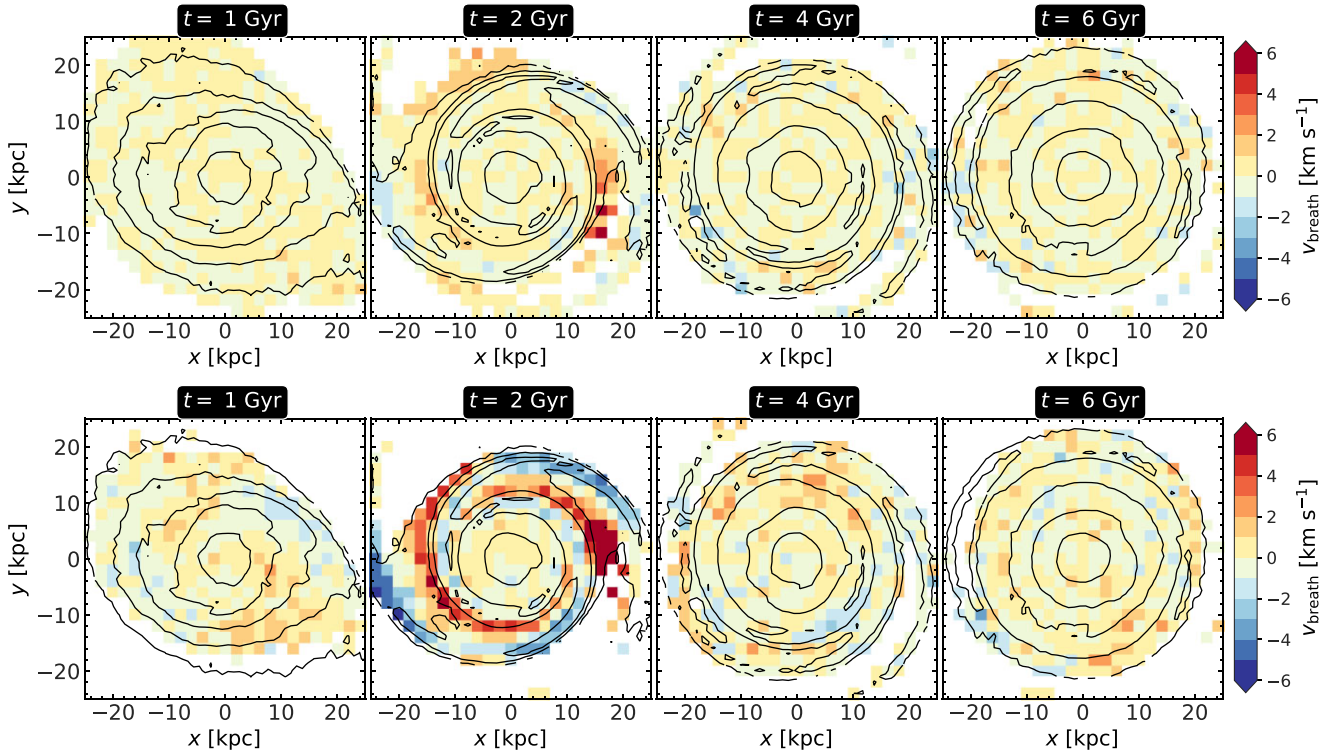
to weak (for details see Section 3.1). At  $t = 1$  Gyr, just after the fly-by encounter, the distribution of the bulk vertical velocity ( $\langle v_z \rangle$ ) in the  $(\phi, z)$ -plane predominantly shows bending motions, i.e. stellar particles on both sides of the mid-plane are moving coherently in the same direction. At this time, a prominent spiral is yet to form in the host galaxy (see Fig. 4). However, by  $t = 2$  Gyr, the host galaxy shows a prominent spiral (see Fig. 4), and the distribution of the bulk vertical velocity ( $\langle v_z \rangle$ ) in the  $(\phi, z)$ -plane changes drastically. Now, the stellar particles on both sides of the mid-plane are moving coherently towards or away from it, indicating vertical breathing motion dominates. The relative sense of the  $\langle v_z \rangle$  varies as a function of the azimuthal angle. However, at  $t = 4$  Gyr when the tidally-induced spirals have weakened substantially, the distribution of the  $\langle v_z \rangle$  is again seen to be dominated by the bending motions. By the end of the simulation ( $t = 6$  Gyr), the spiral has wound up, and the distribution of  $\langle v_z \rangle$  remains dominated by the bending motions of the stars.

To quantify the breathing motions, we define the breathing velocity,  $V_{\text{breath}}$ , as (Debattista 2014; Gaia Collaboration et al. 2018; Ghosh et al. 2022a)

$$V_{\text{breath}}(x, y) = \frac{1}{2}[\langle v_z(x, y, \Delta z) \rangle - \langle v_z(x, y, -\Delta z) \rangle], \quad (10)$$

where  $\langle v_z(x, y, \Delta z) \rangle$  is the mean vertical velocity at position  $(x, y)$  in the galactocentric Cartesian coordinate system, averaged over a vertical layer of thickness  $\Delta z$  (for details see Gaia Collaboration et al. 2018; Ghosh et al. 2022a). A positive breathing velocity ( $V_{\text{breath}} > 0$ ) implies that the stars are coherently moving away from the mid-plane (expanding breathing motion), while  $V_{\text{breath}} < 0$  implies that the stars are moving coherently towards the mid-plane (compressing breathing motion). We find that at larger heights, the particle resolution of our selected fly-by model is not well-suited to compute meaningful values of  $V_{\text{breath}}$ . Therefore, we calculate the distribution of  $V_{\text{breath}}$  in the  $(x, y)$ -plane for two vertical layers, namely,  $|z| = [0, 400]$  pc and  $|z| = [400, 1200]$  pc. The resulting distributions of  $V_{\text{breath}}$  at the same four times (as in Fig. 9) are shown in Fig. 10. The breathing velocity is close to zero near the mid-plane for all four times considered. However, at  $t = 2$  Gyr when the strong spirals are present in the host galaxy, the upper vertical layer ( $|z| = [400, 1200]$  pc) shows significant coherent breathing velocity ( $\sim 6\text{--}8 \text{ km s}^{-1}$ , in magnitude). This trend is similar to what is shown for spiral-driven breathing motions (e.g. see Debattista 2014; Ghosh et al. 2022a), and is also similar to the breathing motions seen from the *Gaia* DR2 (Gaia Collaboration et al. 2018). A visual inspection also reveals that at  $t = 2$  Gyr, the compressing breathing motions are closely associated with the spiral arms whereas the expanding breathing motions arise in the inter-arm regions (see Fig. 10), similar to the findings of Faure et al. (2014), Debattista (2014) and Ghosh et al. (2022a). However, after  $t = 4$  Gyr, when the spirals are either significantly weaker or completely absent, the host galaxy does not show any prominent, coherent breathing motion in the upper layer ( $|z| = [400, 1200]$  pc). To probe this further, we calculate the distribution of  $V_{\text{breath}}$  at two vertical layers in every 200 Myr from  $t = 2$  to 6 Gyr for the model *RP40i00pro*. For the sake of brevity, this is not shown here. We find that the incidence of a prominent, coherent breathing motion is strongly related with the presence of a strong spiral feature in the host galaxy. This together with the amplitude of the breathing velocity increasing with height (a signature of spiral-driven breathing motion, see Ghosh et al. 2022a) indicate that the breathing motions present in the host galaxy are driven by the tidally-induced spirals.

Finally, we study the azimuthal variation of the breathing motions and their connection with the peak(s) of the spirals. Figs 9 and 10 al-

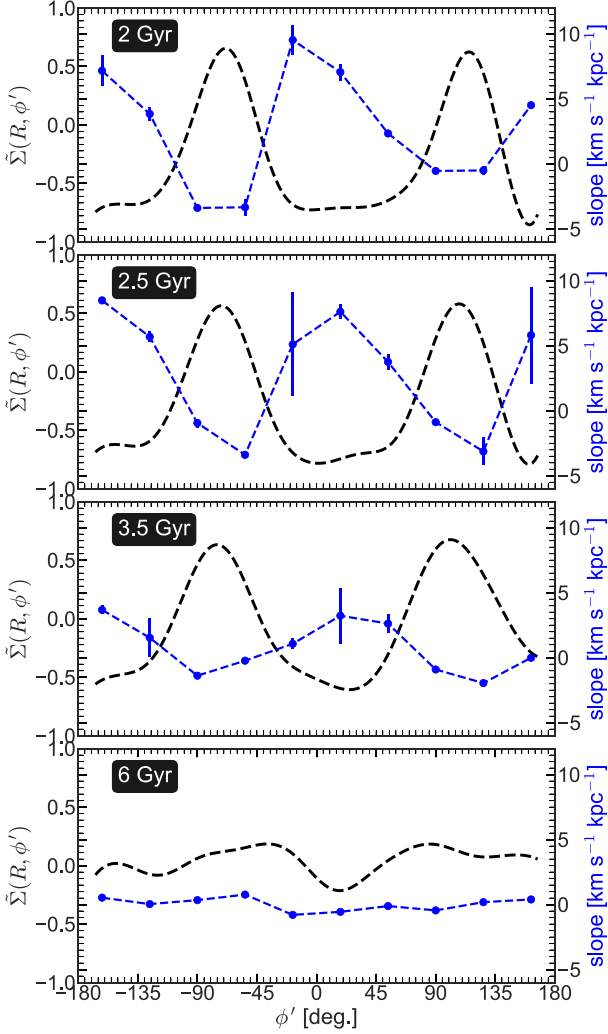


**Figure 10.** Distribution of the breathing velocity,  $V_{\text{breath}}$  (equation 10) at different vertical distances from the mid-plane at four different times for the model *RP40i00pro*. The solid black lines denote the contours of constant density. The *top panels* show stars at  $|z| = [0, 400]$  pc whereas the *bottom panels* show stars at  $|z| = [400, 1200]$  pc, respectively.

ready demonstrated that the breathing motions in model *RP40i00pro* is associated with the spiral arm and the inter-arm regions. Here, we explore this further. Following Ghosh et al. (2022a), we quantify these breathing motions whose amplitude increases with height from the mid-plane by fitting a straight line. In this formalism, the presence of a prominent breathing motion would result in a significantly *non-zero* slope. Furthermore, an expanding (positive) breathing motion will yield a positive slope whereas a compressing (negative) breathing motion will yield a negative slope. On the other hand, the non-zero value of the intercept of the best-fit straight line indicates the presence of a bending motion. For details, the reader is referred to Ghosh et al. (2022a). A similar approach was also used in Widrow et al. (2014). We consider the same radial extent  $15 \text{ kpc} \leq R \leq 17 \text{ kpc}$ , where a prominent spiral is present at  $t = 2 \text{ Gyr}$ . Since the phase-angle of the  $m = 2$ , Fourier mode ( $\phi_2$ ) varies as a function of radius, indicating that the azimuthal locations of the density peaks vary as a function of radius to obtain a stronger signal of the slope, we first rotate the stellar particles in two different 1 kpc-wide radial bins in such a way that the density peaks in our chosen radial extent coincide. Then, we recalculate the slope of the breathing velocity as a function of the rotated azimuthal angle ( $\phi'$ ). The resulting variation of the slope as well as the residual surface density ( $\tilde{\Sigma}(R, \phi')$ ) with rotated azimuthal angle ( $\phi'$ ) are shown in Fig. 11 for four different times. As seen clearly from Fig. 11, at  $t = 2 \text{ Gyr}$  when the spirals are quite strong ( $|A(m = 2, p = p_{\text{max}})| = 0.38$ ), the compressive breathing motions are associated with the peak(s) of the spiral whereas the expanding breathing motions are associated with the density minima. This trend is consistent with the signature of a spiral-driven breathing motions as shown in Debattista

(2014) and Ghosh et al. (2022a). We show that in our model, the amplitudes of the expanding breathing motions are higher than that of the compressing breathing motions – a trend attributed to the fact of having a more abrupt density variation as the stellar particles leave the spirals compared with when they enter the spirals (see Debattista 2014). At  $t = 3.5 \text{ Gyr}$ , when the spirals are decaying ( $|A(m = 2, p = p_{\text{max}})| = 0.21$ ), the corresponding (absolute) values of the slope also decrease, thereby indicating that the breathing motions also weaken. By  $t = 6 \text{ Gyr}$ , the spirals get wound up (almost) completely ( $|A(m = 2, p = p_{\text{max}})| = 0.04$ ), and the (absolute) values of the slope are close to zero, indicating the absence of breathing motion. This further strengthens the case for the breathing motions being driven by the spirals, and not the *direct* dynamical consequence of a fly-by interaction. The breathing motions are seen to last for  $\sim 1.5\text{--}2 \text{ Gyr}$  after their generation for the model *RP40i00pro*.

We repeated this whole set of analyses for all the other models to investigate the breathing motions and their connection with the incidence and strength of spirals present in the model. We find that a strong spiral always drives a prominent breathing motion with expanding breathing motions associated with the inter-arm region whereas the compressing breathing motions are associated with the peak(s) of the spirals. As the spirals get wound up, the breathing motions also cease to exist in the fly-by models. These trends are similar to what we have found for the model *RP40i00pro*. Therefore, for brevity, these are not shown here. In the fly-by model with retrograde orbital configuration (*RP40i00ret*), the spiral structure itself is weaker compared to the other models in prograde orbital configuration (see Fig. 8). We checked that no prominent breathing motion is excited by this feeble spiral structure.



**Figure 11.** Variation of the slopes as a function of the rotated azimuthal angle ( $\phi'$ ) at four different times (having different spiral strength) for the model *RP40i00pro* (see blue dashed lines). Only the stellar particles in the radial extent  $15 \leq R/\text{kpc} \leq 17$  are chosen here. The black dashed line denotes the residual surface density ( $\tilde{\Sigma}(R, \phi')$ ) as a function of the azimuthal angle, calculated in this chosen radial extent. The particles have first been binned in 1 kpc annuli and then azimuthally rotated relative to each other so that the minimum in  $\tilde{\Sigma}(R, \phi')$  in each annulus is coincident, and then the mean vertical velocity ( $\langle v_z \rangle$ ) is calculated, amounting to stacking different radial ranges while unwinding the spiral. For details, see text. The sense of rotation is towards increasing  $\phi'$ .

## 5 DISCUSSION

Our fly-by interactions excite strong spirals in the outer regions of the host galaxy’s disc. The spirals show a variation in their maximum strength depending on the angle of interaction, and the orbital spin vector. Here, we compare the strength, location, and nature of the tidally-induced spirals in our models with past studies. The numerical simulations of Oh, Kim & Lee (2015b) showed that in their models, a stronger tidal encounter induces prominent spirals in the inner regions ( $5 \leq R/\text{kpc} \leq 10$ ) of the host galaxy’s disc (in addition to the tidal tails in the outer parts). The arm strength of the spirals vary in the range  $\sim 0.1\text{--}0.18$ , depending on the values of the relative tidal force, and the relative imparted momentum. Moreover, the spirals are of

kinematic density wave in nature. Also, the simulations of Semczuk et al. (2017) showed the generation of transient spirals due to the tidal force exerted by the potential of a cluster; the spirals appear with each pericentre passage followed by a fast decay. These spirals appear in the outer disc region ( $12 \leq R/\text{kpc} \leq 17$ ) with the maximum arm strength varying in the range  $\sim 0.4\text{--}0.75$ . In comparison, the spirals in our fly-by models are most prominent in the outer regions ( $R \geq 10$  kpc). The maximum arm strength varies in the range  $\sim 0.15\text{--}0.38$ , depending on the angle of interaction, and the orbital spin vector (see Fig. 8). The spirals appear shortly after the pericentre passage of the perturber, grow rapidly, followed by winding up of those spirals. This winding of spirals show two distinct phases, namely, the initial rapid winding phase, followed by a slow winding phase (see Fig. 6).

The amplitude of the spiral-driven breathing motions also merits some discussion. In our fly-by models, when the spirals are most prominent, the values of the best-fit slope (used as a proxy for the breathing amplitude) vary from  $\sim -5$  to  $10 \text{ km s}^{-1} \text{ kpc}^{-1}$  (see Fig. 11). In comparison, the fly-by model (with spirals present) of Widrow et al. (2014) showed the (absolute) values of the best-fit slope  $\sim 10 \text{ km s}^{-1} \text{ kpc}^{-1}$ . As for the spiral-driven breathing motions where the spirals arise due to internal instability, Ghosh et al. (2022a) reported the values of the slope varying from  $\sim -2.5$  to  $3 \text{ km s}^{-1} \text{ kpc}^{-1}$ . Furthermore, the strong spirals present in the models used by Faure et al. (2014), Debattista (2014) can drive large breathing motions ( $|\langle v_z \rangle| \sim 5\text{--}20 \text{ km s}^{-1}$ ).

We have considered only  $N$ -body models of an unbound fly-by interaction excluding the interstellar gas. It is well known that a disc galaxy contains a finite amount of gas (e.g. see Scoville & Sanders 1987). Additionally, in the  $\Lambda$ CDM galaxy formation scenario, a galaxy can accrete cold gas (e.g. Binboim & Dekel 2003; Kereš et al. 2005; Dekel & Binboim 2006; Ocvirk, Pichon & Teyssier 2008), either during the merger-phase or at a later stage. Past studies have shown the dynamical importance of the interstellar gas in the context of cooling the stellar disc and facilitating the generation of fresh spiral waves (Sellwood & Carlberg 1984b), and in maintenance of spiral density waves in infinitesimally-thin discs (Ghosh & Jog 2015, 2016) as well as in a galactic disc with finite thickness (Ghosh & Jog 2022). For gas rich galaxies undergoing such fly-by interactions, the vertical breathing motion may be important for increasing the turbulence in the gas where star formation is insignificant (Stilp et al. 2013). This is because as the stars are in the breathing motion, the vertical potential will change with time and the gas distribution will be affected. However, since the vertical stellar velocity induced by the breathing motion is fairly small (of the order of a few  $\text{km s}^{-1}$ ), this effect may not be very significant, especially if compared to the much larger kinematic effect of supernova explosions and stellar winds (Krumholz et al. 2018; Yu et al. 2021). In addition, galaxy bulges play an important role in maintaining spiral density waves in the disc for a longer time (Saha & Elmegreen 2016). Although, our galaxy models have a classical bulge, we have not varied the contribution of the bulges in our models.

Lastly, our models are specifically designed in a way that the unperturbed disc galaxy is not bar unstable and only forms weak spirals; the strong spirals that form therefore are a dynamical result of the fly-by encounter. However, in reality, the Milky Way also harbours a stellar bar (e.g. see Liszt & Burton 1980; Binney et al. 1991; Blitz & Spergel 1991; Weinberg 1992; Weiland et al. 1994; Dwek et al. 1995; Binney, Gerhard & Spergel 1997; Freudenreich 1998; Hammersley et al. 2000; Wegg & Gerhard 2013). Furthermore, the bulk vertical motions in the Solar neighbourhood and beyond display both bending and breathing motions (e.g. Carrillo et al. 2018; Gaia

Collaboration et al. 2018). This simultaneous presence of bending and breathing motions could well be collectively manifested due to a combination of internal (spiral and/or bar-driven) and external driving mechanisms (tidal encounters), as previously investigated by Carrillo et al. (2019). We stress that the aim of this work is to clarify whether the excitation of breathing motions are ‘directly’ related to tidal interactions or whether they are driven by the tidally-induced spirals (as also mentioned in Section 1), and not to replicate the observed dynamical state of the Milky Way.

## 6 SUMMARY

In summary, we investigated the dynamical impact of an unbound single fly-by interaction with a perturber galaxy on the generation of the tidally-induced spiral features and the associated excitation of vertical breathing motions. We constructed a set of  $N$ -body models of fly-by encounter with mass ratio kept fixed to 5:1 while varying different orbital parameters. Our main findings are:

(i) Fly-by interactions trigger a strong spiral structure in the disc of the host galaxy. The spirals grow rapidly in the initial times, followed by a slow decay. The generation and the strength of these tidally-induced spirals depend strongly on the angle of interaction as well as on the orbital spin vector. For the same orbital energy and the angle of interaction, the models in prograde configuration are more efficient at driving strong spirals when compared to models in retrograde configuration.

(ii) The tidally-induced spirals in the host galaxy can survive for  $\sim 2.9$ – $4.2$  Gyr after their formation. The pitch angle of the resulting spirals display two distinct phases of winding, namely, a fast winding phase ( $\tau_{\text{wind}} \sim 0.4$  Gyr) and a subsequent slow winding phase ( $\tau_{\text{wind}} \sim 5.7$  Gyr).

(iii) When the tidally-induced spirals are strong, they drive coherent, large-scale vertical breathing motions whose amplitude increases with height from the mid-plane. Furthermore, the azimuthal locations of the compressing breathing motions ( $V_{\text{breath}} < 0$ ) are associated with the peaks of the spirals whereas the azimuthal locations of the expanding breathing motions ( $V_{\text{breath}} > 0$ ) coincide with the density minima of the spirals. These trends are in agreement with the signatures of spiral-driven breathing motions.

(iv) The temporal evolution of these breathing motions follow closely the temporal evolution of the strengths of the spirals. A stronger spiral drives breathing motions with larger amplitudes. These breathing motions excited by tidally-induced spirals can persist for  $\sim 1.5$ – $2$  Gyr in the disc of the host galaxy.

Thus, the results presented in this paper demonstrate that a strong spiral structure can drive large coherent vertical breathing motions irrespective of their formation scenario, i.e. whether induced by tidal interactions (as shown here) or generated via internal disc gravitational instability (e.g. Debattista 2014; Faure et al. 2014; Ghosh et al. 2022a). Furthermore, our results highlight the cautionary fact that although in past studies, the tidal interactions are considered as the ‘usual suspect’ for driving the vertical breathing motions, it is indeed the tidally-induced spirals which drive the breathing motions, and the dynamical role of such tidal encounters remains only ancillary.

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## DATA AVAILABILITY

The data generated in this research will be shared on reasonable request to the corresponding author.

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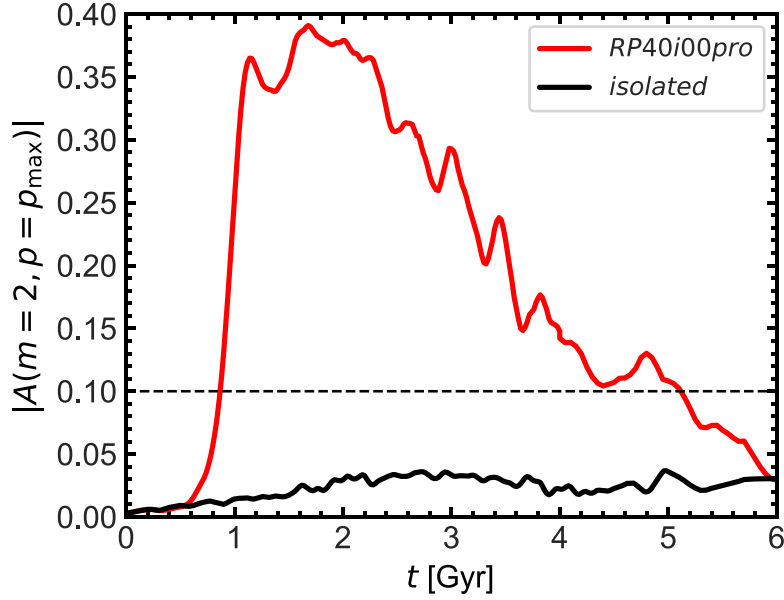
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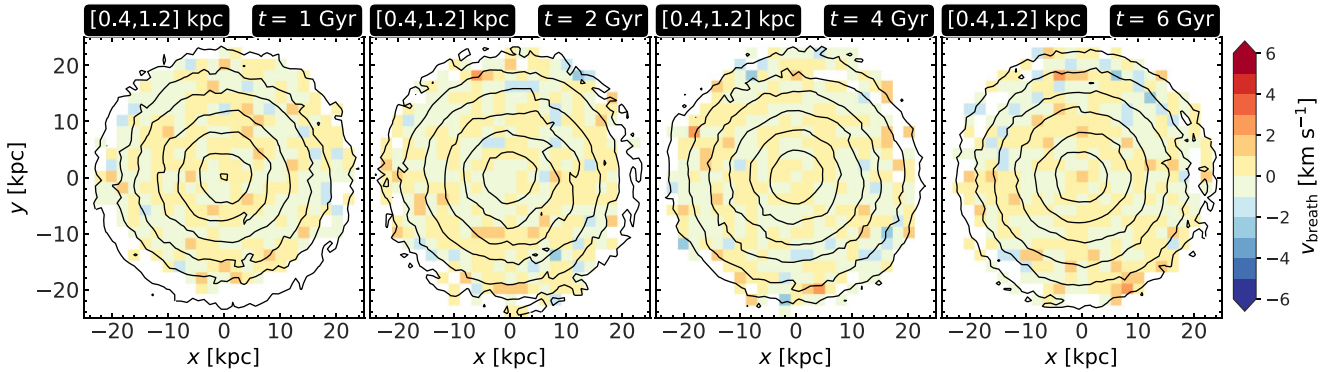
## APPENDIX: EVOLUTION IN ISOLATION

So far, we have shown that in our fly-models, a prominent spiral appears shortly after the interaction happens, and this spiral drives a coherent vertical breathing motion in the host galaxy. However, it remains to be investigated whether the host galaxy, when evolved in isolation, could still generate spirals and the associated vertical breathing motions. We evolve the host galaxy model in isolation for 6 Gyr. A visual inspection of the face-on density distribution of the stellar particles reveals no prominent spirals throughout the simulation run (as can be seen in the density contours in Fig. A2). Following the methodology described in Section 3.1, we measure the values of  $|A(m = 2, p = p_{\max})|$  (used to quantify the strength of spirals) for the isolated host galaxy model. The resulting temporal evolution is shown in Fig. A1. The values of  $|A(m = 2, p = p_{\max})|$  remain close to zero throughout the simulation run demonstrating that no prominent spiral arm is generated during the entire isolated evolution of the host galaxy model.

Furthermore, we calculate the breathing motions ( $V_{\text{breath}}$ ) using equation (10) for both the vertical slices, namely,  $|z| = [0, 400]$  pc and  $|z| = [400, 1200]$  pc. As before, the vertical slice  $|z| = [0, 400]$  pc does not show any breathing motion. Interestingly, the upper vertical slice ( $|z| = [400, 1200]$  pc) does not show any prominent breathing motion either throughout the simulation run (see Fig. A2) in sharp contrast with the fly-by models (compare with Fig. 10). This clearly demonstrates that the spirals in the fly-by models are indeed tidally-induced. In other words, the generation of spirals and the associated spiral-driven vertical breathing motions can indeed be attributed to the dynamical impact of a fly-by interaction.



**Figure A1.** Temporal evolution of the spiral strength ( $|A(m = 2, p = p_{\max})|$ ), calculated using equation (8) is shown for the isolated host galaxy model. For comparison, we kept the model *RP40i00pro* here (see red solid line). The isolated model does not show any prominent spirals throughout the simulation run.



**Figure A2.** Evolution in isolation: distribution of the breathing velocity,  $V_{\text{breath}}$ , for stars at  $|z| = [400, 1200]$  pc at four different times for the isolated host galaxy model. The solid black lines denote the contours of constant density. No prominent vertical breathing motion is seen for the isolated evolution in sharp contrast with the fly-by models (e.g. compare with Fig. 10).

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