# Near-Infrared Observations of RR Lyrae variables in Galactic Globular Clusters: I. The case of M92 

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#### Abstract

We present near-infrared J,H, and K-band time series observations of the Galactic Globular Cluster (GGC) M92. On the basis of these data, we derived well-sampled light curves for eleven out of the seventeen cluster RR Lyrae variables, and in turn, accurate mean near-infrared (NIR) magnitudes. The comparison between predicted and empirical slopes of NIR Period-Luminosity (PL) relations indicates a very good agreement. Cluster distance determinations based on independent theoretical NIR $P L$ relations present uncertainties smaller than $5 \%$ and agree quite well with recent distance estimates based on different distance indicators. We also obtained accurate and deep NIR color-magnitude diagrams (CMDs) ranging from the tip of the Red Giant Branch (RGB) down to the Main Sequence Turn-Off. We detected the RGB bump and the NIR luminosities of this evolutionary feature are, within theoretical and empirical uncertainties, in good agreement with each other.


Subject headings: globular clusters: individual NGC 6341 - stars: evolution stars: variables

## 1. Introduction

RR Lyrae stars are the cornerstone of Population II distance indicators and countless theoretical and observational investigations have been devoted to the intrinsic accuracy of RR Lyrae distance determinations (Caputo et al. 2000; Layden 2002; Cacciari \& Clementini 2003; Walker 2003). However, we are still facing the problem that distance measurements based on different methods present discrepancies larger than the estimated total error budget (see e.g., Carretta et al. 2000; Bono 2003). This might be suggestive of still poorly constrained systematic errors. Current predictions concerning the luminosity of Horizontal Branch (HB) models inside the RR Lyrae instability strip bracket approximately 0.1 dex in luminosity. The difference is due to the adopted input physics which is not well-established yet (Cassisi et al. 1998; Pietrinferni et al. 2004, and references therein).

Notwithstanding this, one of the most widespread methods to estimate RR Lyrae distances to GGCs and Local Group galaxies is the correlation between the visual magnitude and the metal abundance $\left(M_{V} \propto[F e / H]\right)$. This method is in principle very robust, but presents several potential drawbacks: i) a (strong) dependence on evolutionary effects (Demarque et al. 2000; Castellani 2003); ii) a possible nonlinear dependence on the metallicity (Castellani, Chieffi, \& Pulone 1991; Caputo et al. 2000); iii) a non-trivial dependence on the reddening correction and on uncertainties affecting abundance measurements (Dall'Ora et al. 2004). On the other hand, empirical evidence dating back to the 1990s indicates that $K$-band NIR observations of RR Lyrae stars appear marginally affected by deceptive errors. In particular, $K$-band data of RR Lyrae stars, when compared with optical $B, V$-band data, present: i) a smaller dependence on reddening uncertainties; ii) a smaller dependence on metal abundance; iii) a luminosity amplitude which is approximately a factor of 4 smaller than in the $B$-band. This means that with a limited number of phase points, the time-averaged $K$ magnitude can be estimated with very good accuracy, since the light curve in this band is almost sinusoidal. Moreover, and even more importantly, Longmore et al. $(1986,1990)$ demonstrated, on the basis of K-band photometry for a good sample of GGCs, that cluster RR Lyrae do obey to a well-defined $K$-band PL relation $\left(P L_{K}\right)$. This finding has been soundly confirmed by Butler (2003) and by Storm (2004) for RR Lyrae in M3 and in IC4499. The physical basis of the $P L_{K}$ relation has been widely discussed by Bono et al. (1999; 2003). More recently, theoretical investigations (Bono 2003; Catelan, Pritzl, \& Smith 2004) have also emphasized the existence of well-defined PL relations in $J, H$
bands. However, observations lag theoretical predictions, and indeed we still lack accurate mean $K$-band magnitudes of cluster RR Lyrae, since the current sampling of the light curves is rather poor, typically of the order of 4-6 phase points (Longmore et al. 1990). Moreover, $J, H$ measurements are only available for a limited sample of field RR Lyrae stars (Fernley et al. 1987; Carney et al. 1995). In order to fill this observational gap and to validate current predictions concerning the $J, H, K$-band $P L$ relations, we have undertaken a long-term project aimed at improving the current database of NIR observations of cluster RR Lyrae stars (Dall'Ora et al. 2004). We selected a dozen of northern and southern GGCs which cover a wide metallicity range and host a sizable sample of RR Lyrae stars.

In this investigation, we present the results for the GC M92, one of the most metalpoor and oldest clusters in our sample. This cluster has often been a fundamental target to investigate the absolute age of GGCs and the chronology of the Galaxy (Grundahl et al. 2000; VandenBerg 2000; Gratton et al. 1997; Carretta et al. 2000, and references therein). Therefore, accurate and independent distance determinations of this cluster can play a crucial role in the improvement of age estimates (Renzini 1992). In the next section, we discuss the empirical data as well as the reduction strategy adopted to perform the NIR photometry. In $\S, 3$ we present optical and NIR CMDs, together with a detailed discussion concerning the key evolutionary features. In $\S 4$, we discuss NIR light curves for the sample of RR Lyrae stars in M92 and compare empirical and theoretical NIR PL relations, while $\S 5$ deals with the new cluster distance determinations based on the NIR PL relations and the comparison with similar estimates based on various distance indicators. A brief summary and the conclusions are outlined in the final section.

## 2. The observational scenario

### 2.1. The cluster

The GGC M92 (NGC 6341) is very metal-poor ( $[\mathrm{Fe} / \mathrm{H}]=-2.24$, Zinn 1985) and is located at a distance from the Galactic center of $R_{G C}=9.1 \mathrm{kpc}$, and $Z_{G C}=4.3 \mathrm{kpc}$ above the Galactic plane (Harris 1996). Therefore, it is a very good target to investigate its stellar content. The most recent wide-field CCD photometry was obtained in V and I bands with the CFH12K mosaic camera available at the 3.6 m Canada-France-Hawaii Telescope (CFHT) by Lee et al. (2003). The data presented in this paper cover a sky area significantly larger than the tidal radius of M92. The CMD extends from the bright region down to 5 mag fainter than the Main Sequence Turn-Off. NIR photometry ( $J, H, K$ ) of M92 was collected by Davidge \& Courteau (1999), who used the CFHT Adaptive Optics Bonnette for obtaining high angular resolution JHK images of the cluster center. NIR $(J, K)$ observations of giant
stars have been recently provided by Valenti et al. (2004; hereinafter VFPO) using the NIR camera available at the Telescopio Nazionale Galileo (TNG) to investigate evolutionary properties of cluster RGB stars. Finally, optical (WFPC2, Piotto, Cool, \& King 1997) and NIR (NICMOS, Lee et al. 2001) photometry have also been collected with the Hubble Space Telescope.

M92 hosts a sample of 21 variable stars (Sawyer-Hogg 1973; Clement et al. 2001), 17 of them are RR Lyrae stars, 2 are SX Phoenicis, and one is a BL Herculis, i.e. a low-mass evolved HB star (Bono et al. 1997). A sizable sample of 60 bright RG stars, including also three candidate variables detected by Walker (1955), were checked for variability by Welty (1985), but none was confirmed variable. During the last few years, the sample of RR Lyrae stars has been investigated more in the optical than in the NIR bands (Carney et al. 1992; Cohen 1992; Cohen \& Matthews 1992; Storm, Carney \& Latham 1992, hereinafter SCL; Storm et al. 1992; Kopacki 2001; Piersimoni et al. 2003; Tuairisg et al. 2003). In particular, NIR observations for 3 RR Lyrae stars have been collected by Cohen (1992) and by Storm, Carney, \& Latham (1994) to estimate the distance to M92 using the Baade-Wesselink method.

In the following, we adopt the metal-abundance $([F e / H]=-2.25)$ and the $\alpha$-element enhancement $([\alpha / F e]=0.30)$ given by Salaris \& Cassisi (1996, see also Sneden, Pilachowski, \& Kraft 2000). By adopting these values and the global metallicity ( $[M / H]$ ) relation provided by Salaris, Chieffi, \& Straniero (1993), we obtain for M92 a global metallicity $[M / H]=$ -2.04 .

### 2.2. Observations and data reduction

The NIR data for M92 have been collected using the AZT24 1.1 m telescope of the Campo Imperatore (L'Aquila, Italy) observatory, located at an altitude of 2150 m near the Gran Sasso D'Italia mountain (Di Paola 2003). The AZT-24 is equipped with the NIR camera SWIRCAM (Brocato \& Dolci 2003), which is based on a $256 \times 256 \mathrm{HgCdTe}$ "PICNIC" array, sensitive in the wavelength range $1-2.5 \mu m$, provided by Rockwell Science Corp. The camera is provided with the standard wide-band filters $J, H, K, K^{\prime}$ plus additional narrowband filters and grisms. The focal plane scale is $1.04 \mathrm{arcsec} / \mathrm{pix}$, for a total field-of-view larger than $4.4 \times 4.4 \mathrm{arcmin}^{2}$. The FWHM peaks around 2.1 arcsec and typically ranges from 1.8 to 2.8 arcsec , with the best data at 1.6 arcsec. The photometric measurements of M92 were performed in the standard JHK bandpasses. We observed two different fields partially overlapped by $1 \mathrm{x} 1 \mathrm{arcmin}^{2}$ to adopt the same absolute calibration. The fields were chosen to observe the largest number of RR Lyrae stars (11/17). The observations for M92
were carried out during 12 nights between June and September 2002. The typical exposure times range from 2.5 to 7.5 min for the $J$-band, from 2.5 to 10 min for the $H$-band, and from 2.5 to 12.5 min for the $K$-band. The individual exposures were split in a number of sub-exposures ranging from 5 to 15 . For each sub-exposure, we dithered the images by 10 pixels both in X and in Y direction. Target and sky frames were alternatively collected, at a distance of 9 arcmin from each other.

All the images were pre-processed using a pipeline called PREPROCESS, developed by one of us (A. Di Paola). Median sky frames were subtracted from each target image before flat-fielding and co-adding the dithered images. Photometry was performed using DAOPHOT/ALLSTAR (Stetson 1987). In order to improve the photometric accuracy along the light curves we performed several tests using different techniques to stack the images, different criteria to select PSF stars across the individual frames, as well as different reduction strategies to perform the photometry over the entire data set. We found that the best intrinsic accuracy can be achieved by stacking at least 4 consecutive images, due to the low S/N ratio of single images. We selected a reasonable number $(\sim 40)$ of uniformly spaced PSF stars, together with a constant PSF across each frame. To improve the accuracy of individual measurements and the limiting magnitudes, we used a few optical $I$-band images to derive a more complete star catalog ( $\sim 20000$ stars). These images were kindly provided by F . Grundahl (2004, private communication), and have been collected at the Nordic Optical Telescope in La Palma. Finally, we reduced NIR ( $J, H, K$ ) and optical $I$-band frames simultaneously with DAOPHOT/ALLFRAME (Stetson 1994; Dall'Ora et al. 2004).

The standard field of P138-C measured by Persson et al. (1998) was observed 10 times at about the same air mass $(\Delta X<0.02)$ of the observations on M92. In the calibrating equations, we neglected the extinction with the air mass, because the extinction coefficients measured at Campo Imperatore are $<0.1$ for all the $J H K$-bands. We compared the calibrating equations derived for several nights with standard star observations and the accuracy of the NIR calibrations is of the order of 0.03 mag. Current calibrations were also compared with the pioneering work by Cohen, Frogel, \& Persson (1978) for M92. Four stars are in common with their photometric catalogue and the average difference in magnitude is: $\Delta J=0.015 \pm 0.014 \mathrm{mag}, \Delta H=0.004 \pm 0.030 \mathrm{mag}$, and $\Delta K=0.012 \pm 0.013 \mathrm{mag}$ (our magnitudes are brighter). Therefore, the two data sets are, within the errors, fully consistent with no systematic offset.

The comparison between current NIR photometry and $J, K$ data collected by VFPO shows that the two data sets agree quite well, and indeed on a star-by-star basis the difference in magnitude is: $\Delta J=-0.09 \pm 0.05$ and $\Delta K=-0.07 \pm 0.05 \mathrm{mag}$ (with our magnitudes brighter). The large spread is due to the fact that our images and VFPO images overlap
in a small $(1 \times 2$ arcmin $)$ region located close to the cluster center, and we have roughly a dozen of common stars.

## 3. The Color-Magnitude diagrams

The NIR CMDs and the NIR Luminosity Functions (LFs) are powerful diagnostics to investigate RG stars in GCs, since they allow us to constrain the accuracy and reliability of evolutionary predictions concerning low-mass stellar structures (Ferraro et al. 2000, and references therein). Note also that RG stars are the brightest cluster objects in NIR bands, and therefore the unresolved background population marginally affects their photometry. This means that high signal-to-noise NIR measurements of RG stars can be obtained close to the cluster center, i.e. the regions more affected by severe crowding in optical bands. Moreover, since the NIR survey of $\sim 30$ GGCs by Frogel, Cohen, \& Persson (1983), based on a single-channel detector and aperture photometry, it became clear that $V-K$ colors of RG stars are excellent effective temperature indicators. In spite of these indisputable advantages, the number of GGCs for which accurate, and deep NIR CMDs are available, is rather limited.

Figure 1 shows the $K,(J-K)$ and the $H,(J-H)$ (left and middle panel) CMDs of M92, together with the intrinsic photometric errors vs magnitudes (right panels). A glance at the CMDs discloses the sizable sample ( $\approx 1000$ ) of cluster stars which have been measured. It is noteworthy that in the $H,(J-H)$ CMD current photometry approaches the Turn-Off (TO) region and that the intrinsic photometric error is of the order of 0.02-0.03 mag at the magnitude typical of RR Lyrae stars (open and filled circles).

In order to combine visual and NIR magnitudes, we adopted the V-band photometry for M92 provided by Kopacki (2001), because this is the most recent and complete optical investigation of variable stars in M92. This choice offers the advantage of having homogeneous $(V-K)$ colors for both variable and static stars. Figure 2 shows the $J,(V-J), H,(V-H)$ and $K,(V-K)$ CMDs based on the V photometry, kindly provided by Kopacki, and current NIR data. Note that the faint limit of the Kopacki's optical photometry is $\sim 18 \mathrm{mag}$, which is brighter than the turn-off point. However, current NIR photometry in $J$ and $H$ bands attains fainter limiting magnitudes. To overcome this problem, we also cross-correlated the $J, H$ catalogues with the deeper V-band photometry for M92 provided by Rosenberg et al. (1999). Detailed checks on the stars in common in the two V-band catalogues ( $\approx 500$ stars) show that they agree very well, and indeed the mean difference is of the order of $0.02 \pm 0.04$ mag. The main evolutionary features of the optical/NIR CMDs are the following:

- the RGB is well-populated and roughly covers the entire RGB extension. Empirical data range from the RGB tip down to the sub-giant branch $K \approx 17$ mag.
- The sample of RGB stars appears large enough - when compared with the old analysis of Frogel, Cohen, \& Persson (1983) - to provide a reliable estimate of the RGB tip.
- Data plotted in Fig. 2 disclose a well-defined Horizontal Branch (HB) in the three CMDs. Moreover, hot HB stars brighter than $J \approx H \approx K \approx 16$ mag display a welldefined slope when moving from hotter to cooler effective temperatures. This trend is expected and is caused by the strong increase in the $K$-band bolometric correction towards cooler effective temperatures (see Fig. 1 in Bono 2003). Interestingly enough, this empirical evidence brings forward that $K$-band magnitudes of hot HB stars can also be adopted to estimate cluster distances (Cassisi et al. 2005, in preparation).

Fig. 2 also shows a 12 Gyr isochrone for $\mathrm{Z}=0.0001$, and $\mathrm{Y}=0.245^{1}$. The isochrone was plotted by using a cluster reddening of $\mathrm{E}(\mathrm{B}-\mathrm{V})=0.02 \mathrm{mag}$ and the true distance modulus obtained by using the $P L_{J H K}$ relations (see §4.2). Extinction parameters for both optical and NIR bands have been estimated using the standard extinction model of Cardelli et al. (1989). Data plotted in this figure disclose that a cluster age of $\sim 12$ Gyr nicely accounts for stars located around the TO region. This age estimate is, within the errors, in good agreement with age determinations provided by Salaris \& Weiss (2002), Grundahl et al. (2000), and by Carretta et al. (2000) on the basis of deeper and more accurate visual CMDs.

### 3.1. The RGB bump and tip

The RGB Luminosity Function (LF) is a very robust observable in order to check the accuracy of the inner structure of RGB evolutionary models. In particular, the RGB LF can supply tight constraints upon the chemical stratification inside these stars (Renzini \& Fusi Pecci 1988). The hydrogen stratification encountered by the thin H-burning shell affects the evolutionary rate, and in turn, the star counts along the RGB. In particular, the RGB bump marks the evolutionary phase during which the H -burning shell crosses the chemical discontinuity left over by the convective envelope soon after the first dredge-up. Dating back to its first detection in 47 Tuc (Lee 1977), several theoretical and observational

[^0]investigations have been focussed on the RGB bump (Ferraro et al. 1999; Zoccali et al. 1999, and references therein). Owing to the anti-correlation of the RGB bump luminosity with the stellar metallicity, the detection of this feature is much easier in metal-intermediate and in metal-rich GGCs. This is the reason why the RGB bump has only been detected in a few metal-poor GCs. Using both the differential and the cumulative K-band LF (see Fig. 3), we detected the RGB bump at $K_{\text {bump }}=11.90 \pm 0.08 \mathrm{mag}$.

The RGB bump in M92 has also been provided by VFPO, who found a value $K_{b u m p}=$ $12.40 \pm 0.05 \mathrm{mag}$. This estimate is 0.5 mag fainter than current determination. The difference is larger than the estimated intrinsic photometric errors and the measured systematic offset between the two data sets. The difference might be due to the different size of RGB star samples. A glance at the data plotted in fig. 1 of VFPO suggests that their RGB is not very well populated, and thus their estimate of the error affecting the RGB Bump luminosity could have been underestimated. To further constrain the plausibility of the current $K$-band RGB bump detection, we performed a detailed comparison with updated evolutionary predictions (see Fig. 4) provided by Pietrinferni et al. (2004). In order to estimate the absolute K magnitude of the RGB bump in M92, we adopted the true distance modulus based on the RR Lyrae $P L_{J, H, K}$ relations (see $\S 4.2$ ). In order to perform an exhaustive comparison between theory and observations, we also included the RGB bump estimates provided by Ferraro et. (2000) and by VFPO. Data plotted in Fig. 4 indicate that a fairly good agreement does exist between updated theoretical models (solid lines) and observations (see also Cassisi \& Salaris 1997, VFPO, and references therein). However, current $M_{K}($ bump $)$ measurement of M92 appears at variance with the empirical trend (dashed line) found by Valenti et al. (2004a,b). Finally, we note that the observed $K_{b u m p}$ is located at $(V-K)_{0}=2.27$. This means $V_{\text {bump }}=14.17 \mathrm{mag}$ and, by adopting the same true distance modulus, $M_{V}($ bump $) \approx-0.5 \mathrm{mag}$, a value which agrees with theoretical predictions. Stellar isochrones provided by Pietrinferni et al. (2004) predict, for an age of $12 \mathrm{Gyr}, M_{V}($ bump $) \approx-0.42 \mathrm{mag}$.

The current sample of RGB stars appears, in the brighter portion of the RGB, large enough to estimate the brightness of the RGB tip. If we assume that the brightest and coolest RGB star is at the RGB tip, one derives that $K_{t i p}=8.95 \pm 0.2 \mathrm{mag}$. By using the NIR calibration of the RGB tip provided by Bellazzini et al. (2004), i.e. $M_{K}(t i p)=-(0.64 \pm$ $0.12)[M / H]-(6.93 \pm 0.14)$ and by adopting for M92 a global metallicity $[M / H]=-2.04$, we obtain $M_{K}($ tip $)=-5.62 \pm 0.14$ mag. Note that if we adopt the true distance modulus based on the RR Lyrae $P L_{J, H, K}$ relations, we find that the absolute $K$ magnitude of the RGB tip is $M_{K}=-5.67 \pm 0.2$ mag. Therefore, the two independent estimates are, within the errors, identical.

## 4. The RR Lyrae variables

Current NIR observations allowed us to derive accurate light curves for 11 out the 17 RR Lyrae stars and for the BL Her. The light curves in the $J, H, K$ bands for first overtone variables are shown in Figure 5, while the light curves for the sample of fundamental RR Lyrae stars and the BL Her (V7) are displayed in Figure 6. Individual $J, H, K$ measurements are listed in Table 1: for each star we give the identification number, the Heliocentric Julian Day (HJD) of the observations, and the corresponding $J, H, K$ magnitudes together with intrinsic photometric errors. We have determined the mean $J$ and $H$ magnitudes of RR Lyrae by fitting the individual phase points measured by ALLFRAME with a cubic spline. The typical errors on the mean magnitudes are of the order of 0.02 mag . On the other hand, the fit of the individual $K$-band phase points was performed using a template curve. This approach further improves the intrinsic accuracy of mean $K$ magnitudes, because: i) it avoids the spurious fluctuations introduced by the binning of individual phase points; ii) it allows a better propagation of individual photometric errors upon the mean magnitude. This method was originally developed by Jones, Carney \& Fulbright (1996), and provides the mean $K$-band magnitude of RR Lyrae stars by fitting an empirical template light curve to a few phase points. They provided five different templates according to the pulsation mode (fundamental, $R R_{a b}$; first overtone $R R_{c}$ ) and to the $B$-amplitude of the variable. Following this approach and by using the $V$-amplitudes provided by Kopacki (2001), we computed the mean $K$ magnitude for each star by using the template curve. Note that current mean $J, H, K$ magnitudes have been estimated as intensity averages and then transformed into magnitudes (Bono, Caputo, \& Stellingwerf 1995). Moreover, the $V$-amplitudes were scaled into $B$-amplitudes using the empirical relations provided by Jones et al. (1996).

We already mentioned that NIR $J, K$ photometry for the RR Lyrae stars in M92 was only collected by Storm, Carney, \& Latham (1992, hereinafter SCL) and by Cohen \& Matthews (1992, hereinafter CM). For the sake of comparison, Table 2 gives mean $K$ and $J$ magnitudes for the variables in common with the quoted authors. The agreement, for variables V1 and V3, among independent measurements is better than 0.02 mag . For variable V6, the agreement with CM is better than 0.01 in the $K$-band, but we find a discrepancy of $\sim 0.08$ mag in the $J$-band. This discrepancy has no immediate explanation.

The pulsation parameters for RR Lyrae stars in our sample together with their intensity averaged mean $J, H, K, V$ magnitudes and luminosity amplitudes are listed in Table 3. The periods, the mean $\langle V>$ magnitudes, and the $V$-band amplitudes are from Kopacki (2001). A few variables in our sample require individual comments. The variable V4 has a crowded neighborhood, and indeed the optical $I$-band images show a neighbor star located at a distance $\sim 1$ arcsec from the variable, i.e. a distance similar to the NIR image resolution.

Therefore, its pulsational properties are unreliable, and in the following analysis we excluded this variable from our sample. The variable V7 was classified by Kopacki (2001), at variance with the Sawyer-Hogg catalogue (1973), as a BL Her instead of a fundamental RR Lyrae. Our NIR mean magnitudes confirm this cassification, since this variable is on average one mag brighter than typical RR Lyrae stars.

### 4.1. Stellar distribution along the Horizontal Branch

On the basis of an accurate optical photometry, Carney et al. (1992) noted that a static star, in particular a standard star (listed by Sandage \& Walker 1966), was located inside the instability strip of M92. This occurrence is not unique, and indeed Silbermann \& Smith (1995) found a similar evidence in the instability strip of M15. The explanation suggested by Carney et al. (1992) was that $B-V$ colors are not good effective temperature indicators. To further investigate this point, Figure 7 shows the distribution in the $(V-K), K$ plane of both RR Lyrae and static HB stars ( $\times$ symbols) in M92. The predicted edges of the instability strip are plotted as solid lines and are based on RR Lyrae pulsation models provided by Bono et al. (2003). These models were constructed by adopting metal and Helium abundance of $\mathrm{Z}=0.0001, \mathrm{Y}=0.24$, respectively. Predictions were transformed into the observational plane using the atmosphere models provided by Castelli \& Kurucz (2004). In the comparison between theory and observations, we adopted the same reddening and true distance modulus adopted in the isochrone fit. Data plotted in this figure show the expected ranking of the $V-K$ colors with the pulsation period, and indeed first overtone and fundamental pulsators do not overlap. Moreover, we also found three outliers in the instability strip. A first overtone is located outside the instability strip, while two static stars are located inside the instability strip: one is close to the first overtone blue (hot) edge, and one is near the fundamental red (cool) edge. Interestingly enough, the latter static $\operatorname{star}(V-K=1.52, K=13.50)$ is the star identified by Carney et al. (1992). However, these objects in the $(V-K), K$ plane do not pose a severe problem, since their location can be safely explained if we account for theoretical uncertainties affecting both pulsation and atmosphere models (Bono et al. 1997) as well as for empirical uncertainties (see error bars). This finding strongly supports the hypothesis suggested by Carney et al. (1992) to explain the occurrence of static stars inside the instability strip. Finally, we would like to mention that the two static stars located below/above the instability strip might be either photometric blends or field stars.

### 4.2. The NIR period-luminosity relations

During the last few years, a substantial theoretical effort has been devoted to the pulsational properties of RR Lyrae stars in NIR bands. These investigations rely on predictions based either on pulsational models (Bono et al. 2001,2002,2003) or on synthetic horizontal branches (Catelan et al. 2004; Cassisi et al. 2004), and provide PL relations in different NIR bands. More in detail, Cassisi et al. (2004) investigated the predicted behavior of RR Lyrae stars as a function of the HB morphology over a wide metallicity range, namely ( $0.0001<Z<0.006$ ). Interestingly enough, they found that for a $\mathrm{HB} \operatorname{type}^{2}=0.90$, i.e. the HB type of M92, the $\mathrm{PL}_{\mathrm{K}}$ relation does not depend on the metallicity and is the following:

$$
\begin{equation*}
<M_{K}>=-2.30( \pm 0.01)(\log P+0.30)-0.46( \pm 0.01) \tag{1}
\end{equation*}
$$

On the other hand, Catelan et al. (2004) using a similar approach and synthetic HB models covering a narrower metallicity range $0.0005<Z<0.006$, found very similar NIR PL relations. In particular, their relation, for HB Type $=0.934$ and $\mathrm{Z}=0.0005$, are the following:

$$
\begin{align*}
<M_{J}> & =-1.902( \pm 0.045) \log P-0.826( \pm 0.012)  \tag{2a}\\
<M_{H}> & =-2.311( \pm 0.013) \log P-1.136( \pm 0.003)  \tag{2b}\\
<M_{K}> & =-2.343( \pm 0.012) \log P-1.168( \pm 0.002) \tag{2c}
\end{align*}
$$

The two sets of PL relations, for $\mathrm{P}=0.5$ days, $\mathrm{Z}=0.0001$, and HB Type $=0.95$ supply $K$ magnitudes that differ by only 0.01 mag . Therefore, we decide to use the NIR PL relations provided Catelan et al. (2004) to fit our data, though they do not cover the global metallicity of M92.

In order to provide two independent theoretical frameworks to be compared with empirical data we also computed, using the same evolutionary and pulsation models adopted by Cassisi et al. (2004), new $J, H$-band PL relations. We found:

$$
\begin{align*}
< & M_{J}>=-1.708( \pm 0.006)(\log P+0.30)-0.240( \pm 0.004)  \tag{3}\\
& <M_{H}>=-2.26( \pm 0.01)(\log P+0.30)-0.44( \pm 0.01) \tag{4}
\end{align*}
$$

[^1]From an empirical point of view, there are several $P L_{K}$ relations available in the literature, but a general consensus on the slope and on the zero-point has not been reached yet. They range from $M_{K}=-1.92 \times \log P-0.74$ (Longmore et al. 1990), to $M_{K}=$ $-2.95( \pm 0.10) \times \log P-1.07( \pm 0.10)$ (Skillen et al. 1993). In Figure 8, we plot the mean $J, H$, and $K$ magnitudes for our sample of RR Lyrae variables as a function of the fundamentalized period $\left(\log P_{F}=\log P_{F O}+0.127\right)$. The best fits to empirical data ${ }^{3}$ provide the following relations:

$$
\begin{align*}
<J> & =-1.73( \pm 0.24) \log P+13.863( \pm 0.049)  \tag{5}\\
<H> & =-1.85( \pm 0.25) \log P+13.590( \pm 0.051)  \tag{6}\\
<K> & =-2.26( \pm 0.20) \log P+13.475( \pm 0.040) \tag{7}
\end{align*}
$$

In each panel, the empirical best fits (dotted lines) have been overlapped to the data together with the predicted $K$-band PL relation by Cassisi et al. (2004), the new $J, H$ band PL (see equations $1,3,4$; solid lines) and the NIR PL relations provided by Catelan et al. (2004, see equations $2 \mathrm{a}, 2 \mathrm{~b}, 2 \mathrm{c}$; dashed lines). Data plotted in this figure disclose that predicted and observed slopes agree quite well in the three NIR photometric bands, thus supporting the plausibility and the accuracy of current pulsation and evolutionary models.

## 5. Discussion

Current findings together with similar results for cluster and field RR Lyrae stars (Butler 2003; Bono et al. 2003; Borissova et al. 2004; Dall'Ora et al. 2004) support the evidence that NIR PL relations can supply accurate distance estimates (Bono et al. 2001). Therefore, by using the new $J$-band PL relation (eq. 3) and individual mean $J$ magnitudes, we estimated the apparent distance modulus $(m-M)_{J}=<J>-M_{J}$ for each RR Lyrae in our sample. The weighted average of these estimates gives $(m-M)_{J}=14.62 \pm 0.05$ mag. We adopted the same approach for $H$ (eq. 4) and $K$-band measurements and we found $(m-M)_{H}=$ $14.61 \pm 0.06$ and $(m-M)_{K}=14.62 \pm 0.04 \mathrm{mag}$, respectively. The quoted errors account only for the dispersion around the mean of the individual estimates. The typical reddening value for M92 is $E(B-V)=0.02 \mathrm{mag}$ (Harris 1996). By using the standard extinction model by Cardelli et al.(1989), we found that $A_{J}=0.282 \times A_{V}, A_{H}=0.190 \times A_{V}$, and

[^2]$A_{K}=0.114 \times A_{V}$, which give absorptions smaller than 0.02 mag in NIR bands $\left(A_{J}=0.017\right.$, $A_{H}=0.012, A_{K}=0.007$ ).

The final error budget was evaluated by accounting for the dispersion around the mean of the individual estimates, for the absolute photometric calibration $(\approx 0.03)$, and for the uncertainty affecting the zero-point of the predicted PL relation $(\approx 0.01)$. By adding in quadrature these error sources, we obtain a "global" uncertainty of 0.05 mag for the $K$ band, 0.06 , and 0.07 mag for the $J$ and $H$ band, respectively. By accounting for reddening corrections in the three different bands, the true cluster distance modulus, obtained as a weighted average of $J, H$, and $K$ estimates (see Table 4), is equal to $(m-M)_{0}=14.61 \pm 0.03$ $\mathrm{mag}^{4}$. Table 4 also gives the true distance moduli, based on the NIR PL relations predicted by Catelan et al. (2004) ${ }^{5}$. It is worth noting that the two sets of independent distance determinations are, within the uncertainties, identical.

During the last few years, the discussion concerning pros and cons of different methods to estimate the distances to GGCs has been addressed in several theoretical and empirical investigations (Reid 1997; Pont et al. 1998; Carretta et al. 2000; Caputo et al. 2000; Cacciari 2003). In particular, Pont et al. (1998) and Carretta et al. (2000) using the subdwarf Main Sequence fitting found true distance moduli for M92 which agree quite well with each other, $14.61 \pm 0.05 \mathrm{mag}$ and $14.64 \pm 0.07 \mathrm{mag}$, respectively. To further constrain the accuracy of distance determinations to M92, we also adopted the First Overtone Blue Edge (FOBE) method suggested by Caputo et al. (2000). This method relies on the comparison in the Period- $V$-band magnitude plane between predicted and empirical $V$-band first overtone blue edge. By using this method and the visual magnitudes provided by Kopacki (2001) for the three $R R c$ variables in M92, we derived a true distance modulus equal to $14.62 \pm 0.07$ mag. To provide a comprehensive scenario concerning distance determinations to M92, we also used the predicted HB luminosity provided by Pietrinferni et al. (2004). Theoretical HB models predict the luminosity of the Zero Age Horizontal Branch at the level of the RR Lyrae instability strip, i.e. $\log T_{e} \sim 3.85$. However, HB stars in NIR bands do show a well-defined slope. In order to overcome this problem, we used predicted and empirical $V-K$ colors to estimate individual effective temperatures. We selected two variables, V10 and V11, for which we found $\log T_{e}=3.84$ and 3.86 respectively. The mean K magnitude of these two variables is: $<K>=14.24 \pm 0.1$ mag. The predicted absolute $K$-band magnitude of the ZAHB at the metallicity of M92 is equal to: $M_{K}=-0.393 \mathrm{mag}$. Therefore, the true distance

[^3]modulus is equal to $(m-M)_{0}=14.62 \pm 0.1 \mathrm{mag}$. Note that the comparison between the $K$-band mean magnitude of RR Lyrae stars and predicted $K$-band ZAHB brightness at $\log T_{e}=3.85$ appears as a very promising approach, since in this band the evolutionary effects also imply a change in effective temperature (Bono et al. 2001). The seventh column of Table 4 gives the distance to M92 estimated using the near-infrared Baade-Wesselink (BW) method for two cluster RR Lyrae stars observed by Cohen (1992) and by Storm et al. (1994). The last column in Table 4 lists the distance modulus to M92 using the empirical calibration of the luminosity of the RGB tip recently provided by Bellazzini et al. (2004) and shortly discussed in Sect. 3.1

Interestingly enough, data listed in Table 5 disclose that distance determinations based on different methods and photometric bands agree quite well, within current empirical and theoretical uncertainty. This outcome applies to distance estimates based on both empirical and theoretical calibrations of different distance indicators. In particular, the agreement is better than $5 \%$ in distance estimates based on pulsational and evolutionary methods (predicted NIR PL relations, FOBE, ZAHB) and on the MS fitting. An uncertainty of the order of $10 \%$ for the tip of the RGB is expected due to the limited number of stars across this evolutionary phase in metal-poor clusters. Distances based on the BW method present larger uncertainties. The difference might be due to the fact that BW distances are only based on two cluster variables.

## 6. Summary and final remarks

We collected new multiband NIR ( $J, H, K$ ) time series data of the GGC M92. The main results of this investigation can be summarized as follows:

- We performed an extensive and accurate NIR photometric investigation of RR Lyrae stars in the GGC M92. For the first time, we provided well-sampled $J, H, K$-band light curves for eleven cluster RR Lyrae variables.
- We provided accurate estimates of NIR pulsation properties (mean magnitudes, luminosity amplitudes). The comparison between current data and updated theoretical prescriptions discloses a very good agreement concerning the slope of the $J, H, K$-band PL relations.
- By adopting mean $J, H$ and $K$ magnitudes for our sample of RR Lyrae stars and by relying on theoretical PL relations derived in this work for the $J$ and $H$-bands and predicted by Cassisi et al. (2004) for the $K$ band, we estimated the distance to

M92 with an accuracy better than $5 \%$. The same outcome applies if we use the NIR PL relations provided by Catelan et al. (2004). Our distance estimates are also in very good agreement with recent accurate and independent measurements based on various distance indicators. This finding further strengthens the use of predicted NIR PL relations for estimating distances to GGCs and Local Group galaxies.

The current investigation clearly indicates the key-role of new and accurate NIR photometry of RR Lyrae stars in GGCs in constraining their pulsational and evolutionary properties and to assess the reliability and accuracy of NIR PL relations as distance indicators for field and cluster RR Lyrae stars. Needless to say, it is quite important to extend such photometric analysis to other GGCs, in order to investigate the dependence on the metal content, on the evolutionary status, and on HB morphology (Cassisi et al. 2004; Catelan et al. 2004).

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Fig. 1.- NIR Color-Magnitude Diagrams of M92 in ( $K, J-K$ ) (left panel), and in ( $H, J-H$ ) (middle panel). Different symbols mark the location of variable stars in our sample. The right panels show from top to bottom the intrinsic photometric errors in $J, H, K$ measurements as a function of the magnitude. See text for more details.


Fig. 2.- Optical and NIR CMDs of M92 in $(J, V-J)$ (left panel), in $(H, V-H)$ (middle panel), and in ( $K, V-K$ ) (right panel). The solid line shows a 12 Gyr isochrone. The adopted true distance modulus and cluster reddening are $(m-M)_{0}=14.61 \mathrm{mag}$ and $E(B-V)=0.02$ mag as labeled in the figure.


Fig. 3.- The cumulative $K$-band RGB LF. The arrow marks the location of the RGB bump.


Fig. 4.- Comparison between predicted absolute $K$ magnitude of the RGB bump as a function of the global metallicity (Pietrinferni et al. 2004) and observational measurements (Ferraro et al. 2000; VFPO). The diamond marks the location of the RGB bump in M92. The different solid lines display predictions for different assumptions on the cluster age: 12 Gyr (top line), 14 Gyrs (middle line), and 16 Gyr (bottom line). The dashed line is the fit to the empirical measurements provided by Valenti et al. (2004b).


Fig. 5.- $J$ (left), $H$ (center) $K$-band (right) light curves for first overtone RR Lyrae stars in M92. The labels in the left panels show the variable identification and the pulsation period (days). The solid lines in the right panels display the adopted K-band template curve. See text for more details.


Fig. 6.- Same as Fig. 5, but for fundamental RR Lyrae stars and the BL Her variable (V7, bottom panels).


Fig. 7.- Optical NIR CMD for both RR Lyrae stars and static HB stars. The symbols are the same as in Fig. 1. Solid lines show the position of the theoretical instability strip (Bono et al. 2003). The error bars in the top left corner only accounts for observational intrinsic errors.


Fig. 8.- NIR PL relations for RR Lyrae stars in M92, in $J$ (top), $H$ (middle), and $K$ (bottom) bands. The periods of first overtone pulsators ( $R R_{c}$, open circles) have been fundamentalized, i.e. we added 0.127 to $\log P_{F O}$. Filled circles are fundamental RR Lyrae $\left(R R_{a b}\right)$ stars. The dotted lines display fits to empirical data, while the solid lines the theoretical slopes provided by Cassisi et al. (2004) for the $K$-band, and by current investigation for the $J, H$-bands (see eqs. 1,3,4). The dashed lines show the slopes predicted by Catelan et al. (2004, see eqs. 2a, 2b, 2c). The variables V4 and V9 were not included in the empirical fit. The error bar plotted in the middle panel shows the standard deviation of empirical best fits. The standard deviation of theoretical relations are smaller than symbol size.

Table 1: $J, H, K$-band photometry of variable stars.

|  |  | V1 |  |  |  |  |  |  |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
| HJD-2452000 $^{\mathrm{a}}$ | J | $\sigma_{J}{ }^{\mathrm{b}}$ | HJD-2452000 $^{\mathrm{a}}$ | H | $\sigma_{H}{ }^{\mathrm{b}}$ | HJD-2452000 | K | $\sigma_{K}{ }^{\mathrm{b}}$ |
| 438.5094 | 13.919 | .023 | 438.5627 | 13.857 | .029 | 431.9563 | 13.844 | .075 |
| 438.5204 | 13.959 | .022 | 446.5131 | 13.869 | .080 | 439.0346 | 13.763 | .054 |
| 446.4588 | 14.030 | .089 | 460.4153 | 13.794 | .019 | 446.9874 | 13.747 | .147 |
| 460.3892 | 13.981 | .015 | 460.5087 | 13.801 | .020 | 460.9511 | 13.667 | .051 |
| 460.4740 | 14.022 | .019 | 466.9004 | 13.917 | .036 | 465.9515 | 13.829 | .060 |

Note. - Table 1 is published in its entirety in the electronic edition of the Astronomical Journal. A portion is shown here for guidance regarding its form and content.
${ }^{a}$ Heliocentric Julian Day.
${ }^{b}$ Intrinsic photometric errors.

Table 2: Mean $K$ and $J$ magnitudes for the RR Lyrae variables in common with Storm, Carney \& Latham (1992, SCL) and with Cohen \& Matthews (1992, CM).

| ID | $\langle K\rangle_{\text {this work }}$ | $\langle K\rangle_{S C L}$ | $\langle K\rangle_{C M}$ | $\langle J\rangle_{\text {this work }}$ | $\langle J\rangle_{C M}$ |
| :---: | :---: | :---: | :---: | :---: | :---: |
| V1 | 13.81 | 13.823 | 13.800 | 14.07 | 14.066 |
| V3 | 13.96 | 13.958 | $\ldots$ | $\ldots$ | $\ldots$ |
| V6 | 13.91 | $\ldots$ | 13.897 | 14.23 | 14.149 |

Table 3: Pulsational properties of RR Lyrae stars in our sample. The identification numbers are from the Sawyer-Hogg (1973) catalogue, while periods, mean V magnitudes and V-band amplitudes are from Kopacki (2001). We estimated for the mean magnitudes a global error of 0.03 mag , which is dominated by the intrinsic error affecting the absolute zero-point calibration.

| ID | $\begin{aligned} & P \\ & \text { (days) } \end{aligned}$ | $\begin{gathered} \hline\left\langle J>^{a}\right. \\ (\mathrm{mag}) \end{gathered}$ | $\begin{aligned} & \hline \hline A_{J}{ }^{\mathrm{b}} \\ & (\mathrm{mag}) \end{aligned}$ | $\begin{gathered} \hline \hline H>^{\mathrm{a}} \\ (\mathrm{mag}) \end{gathered}$ | $\begin{aligned} & \hline \hline A_{H}{ }^{\mathrm{b}} \\ & (\mathrm{mag}) \end{aligned}$ | $\begin{gathered} \hline \hline K>^{\mathrm{a}} \\ (\mathrm{mag}) \end{gathered}$ | $\begin{aligned} & \hline \hline A_{K}{ }^{\mathrm{b}} \\ & (\mathrm{mag}) \end{aligned}$ | $\begin{gathered} \hline \hline<V>-<K> \\ (\mathrm{mag}) \end{gathered}$ | $\begin{aligned} & \hline \hline A_{V}{ }^{\mathrm{b}} \\ & (\mathrm{mag}) \end{aligned}$ | Type |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
| V1 | 0.7027968 | 14.07 | 0.31 | 13.87 | 0.18 | 13.81 | 0.28 | 1.21 | 0.86 | ab |
| V2 | 0.6438844 | 14.18 | 0.37 | 13.96 | 0.27 | 13.89 | 0.27 | 1.19 | 0.85 | $a b$ |
| V3 | 0.63768 | 14.27 | 0.41 | 13.98 | 0.14 | 13.96 | 0.31 | 1.13 | 1.17 | ab |
| V4 | 0.6289301 | 13.96 | 0.38 | 13.72 | 0.23 | 13.67 | 0.29 | 1.45 | 0.84 | ab |
| V5 | 0.6196771 | 14.30 | 0.34 | 14.04 | 0.10 | 14.01 | 0.30 | 1.11 | 0.98 | ab |
| V6 | 0.6000029 | 14.23 | 0.39 | 13.89 | 0.26 | 13.91 | 0.32 | 1.05 | 1.09 | ab |
| V7 | 1.0614007 | 13.00 | 0.11 | 12.71 | 0.19 | 12.64 | 0.16 | 1.51 | 0.64 | Bl Her |
| V8 | 0.67278 | 14.17 | 0.31 | 13.95 | 0.22 | 13.86 | 0.26 | 1.19 | 0.67 | ab |
| V9 | 0.6085360 | 14.14 | 0.37 | 13.81 | 0.28 | 13.82 | 0.36 | 1.31 | 1.19 | ab |
| V10 | 0.3773043 | 14.38 | 0.23 | 14.18 | 0.10 | 14.14 | 0.11 | 0.97 | 0.48 | c |
| V11 | 0.308445 | 14.53 | 0.18 | 14.30 | 0.16 | 14.35 | 0.11 | 0.75 | 0.62 | c |
| V12 | 0.4099940 | 14.27 | 0.18 | 14.03 | 0.10 | 14.03 | 0.11 | 1.03 | 0.41 | c |

[^4]Table 4: True distance moduli for M92 estimated by using different distance indicators.

| $P L_{J}$ | $P L_{H}$ | $P L_{K}$ | $F O B E$ | MS fitting | ZAHB | BW | RGB Tip |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
| $14.61 \pm 0.06^{\mathrm{a}}$ | $14.60 \pm 0.07^{\mathrm{a}}$ | $14.61 \pm 0.05^{\mathrm{a}}$ | $14.62 \pm 0.07^{\mathrm{c}}$ | $14.61 \pm 0.05^{\mathrm{d}}$ | $14.62 \pm 0.10^{\mathrm{f}}$ | $14.47 \pm 0.15^{\mathrm{g}}$ | $14.56 \pm 0.20^{\mathrm{i}}$ |
| $14.63 \pm 0.05^{\mathrm{b}}$ | $14.61 \pm 0.06^{\mathrm{b}}$ | $14.62 \pm 0.04^{\mathrm{b}}$ | $\ldots$ | $\ldots$ | $14.64 \pm 0.07^{\mathrm{e}}$ | $\ldots$ | $14.60 \pm 0.26^{\mathrm{h}}$ |

${ }^{a}$ Distance determinations based on predicted RR Lyrae $J, H, K$-band PL relations (see eqs. 1,3,4).
${ }^{b}$ Distance determinations based on predicted RR Lyrae $J, H, K$-band PL relations provided by Catelan et al. (2004, see eqs. 2a, 2b, 2c).
${ }^{c}$ Distance determination based on the First Overtone Blue Edge method (Caputo et al. 2000).
${ }^{d}$ Distance determination based on the Main Sequence fitting (Pont et al. 1998).
${ }^{e}$ Distance determination based on the Main Sequence fitting (Carretta et al. 2000).
${ }^{f}$ Distance determination based on ZAHB models (Pietrinferni et al. 2004).
${ }^{g}$ Distance determination based on the BW method (Cohen 1992).
${ }^{h}$ Distance determination based on the BW method (Storm et al. 1994).
${ }^{i}$ Distance determination based on the RGB Tip (Bellazzini et al. 2004).


[^0]:    ${ }^{1}$ The isochrone was retrieved from the WEB site http://www.te.astro.it/BASTI/index.php and is based on updated evolutionary models recently provided by Pietrinferni et al. (2004).

[^1]:    ${ }^{2}$ This parameter was introduced by $\operatorname{Zinn}$ (1980) and is the ratio $N_{B} /\left(N_{B}+N_{R R}+N_{R}\right)$, where $N_{B}$ and $N_{R}$ are the number of HB stars bluer/redder than RR Lyrae stars and $N_{R R}$ is the number of RR Lyrae stars. GGCs characterized by a very blue HBs have HB type close to one, while GGCs with very red HBs have HB type close to zero.

[^2]:    ${ }^{3}$ Note that we did not include variable V9 in the best fit estimates, because it is located near the center of the cluster. Therefore, the pulsational properties of this variable are less accurate than for the other variables. However, if we account for this variable the best fits become: $\langle J\rangle=-1.77( \pm 0.28) \log P+13.845( \pm 0.057)$, $<H>=-1.91( \pm 0.37) \log P+13.558( \pm 0.076)$, and $<K>=-2.30( \pm 0.30) \log P+13.449( \pm 0.061)$.

[^3]:    ${ }^{4}$ This distance estimate does not include the variable V9. If we account for this variable the true cluster distance is $(m-M)_{0}=14.59 \pm 0.04 \mathrm{mag}$.
    ${ }^{5}$ The variable V9 was not included.

[^4]:    ${ }^{a}$ Intensity averaged mean $J, H, K$-magnitudes.
    ${ }^{b}$ uminosity amplitudes in $J, H, K, V$-bands.

