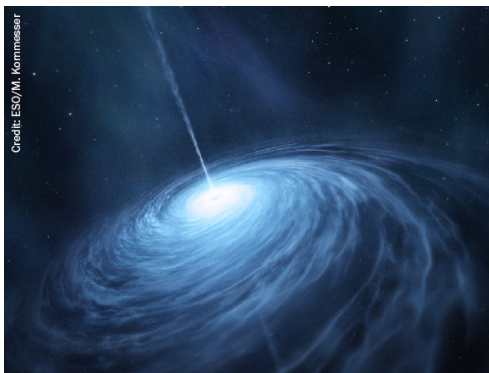


Measurement of the cosmic ray muon flux seasonal variation with the OPERA detector

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Measurement of the cosmic ray muon flux seasonal variation with the OPERA detector



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Abstract. The OPERA experiment discovered muon neutrino into tau neutrino oscillations in appearance mode, detecting tau leptons by means of nuclear emulsion films. The apparatus was also endowed with electronic detectors with tracking capability, such as scintillator strips and resistive plate chambers. Because of its location in the underground Gran Sasso laboratory, under 3800 m.w.e., the OPERA detector has also been used as an observatory for TeV muons produced by cosmic rays in the atmosphere. In this paper the measurement of the single muon flux modulation and its correlation with the seasonal cycle of atmospheric temperature is reported.

Keywords: cosmic rays detectors, neutrino detectors, neutrino experiments

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Contents

1	Introduction	1
2	Cosmic ray muon flux measurement and its modulation	2
3	Atmospheric temperature modulation	4
4	Cosmic ray flux and effective atmospheric temperature correlation	5
5	Conclusions	7
A	Models for modulation of cosmic rays and effective atmospheric temperature	8

1 Introduction

Muons observed in underground laboratories are produced mainly in decays of pions and kaons generated by the interaction of primary cosmic rays with the upper atmosphere. Since muons lose energy crossing the rock overburden, only high energy muons can be detected, with an energy threshold depending on depth, usually expressed in metre water equivalent (m.w.e.). It is known that the flux of atmospheric muons detected deep underground shows seasonal time variations correlated with the temperature of the stratosphere, where the primary cosmic rays interact [1]. This effect has been reported by other experiments at the underground Gran Sasso Laboratory (LNGS) (MACRO [2–4], LVD [5, 6], Borexino [7, 8] and GERDA [9]) and elsewhere (AMANDA [10], IceCube [11], MINOS [12, 13], Double Chooz [14] and Daya Bay [15]). An increase in temperature of the stratosphere causes a decrease of the air density, thus reducing the chance of π, K mesons to interact, and resulting in a larger fraction decaying into muons. Consequently, the atmospheric muon rate changes during the year increasing in summer and decreasing in winter. The variation can be modelled as a sinusoidal function, which is only a first order approximation since the average temperature is not precisely constant over the years and short term effects occur, leading to local maxima and minima, such as “sudden stratospheric warming” events [16]. Additional spectral characteristics can be recovered by autoregressive models (e.g. Box and Jenkins [17]), but in this paper a pure seasonal model is used for consistency with other experimental analyses [2, 15]. Moreover, the primary aim of the present paper, and of those cited, is not the careful description of time evolution, but the study of the correlation between atmospheric temperature and muon flux relative variations, related to pion and kaon production in primary cosmic ray interactions.

The OPERA experiment [18] discovered ν_μ into ν_τ oscillations in appearance mode using the CERN Neutrino to Gran Sasso (CNGS) beam [19, 20]. The experiment was located in the Hall C of LNGS laboratory, at 3800 m.w.e. depth. To identify ν_τ Charged Current interactions, the Emulsion Cloud Chamber technique was used, with 1 mm lead sheets with alternating nuclear emulsion films, for a total target mass of about 1.2 kt. Lead and emulsions were organised into units called “bricks”. The detector was divided into two identical Super-Modules (SM), each made of a target section, composed of 31 brick walls interleaved with

target tracker layers to locate bricks with neutrino interactions, and of a muon spectrometer to optimise the muon identification probability and to measure momentum and charge. In each SM, the Target Tracker (TT) [21] consisted of 31 pairs of orthogonal planes made of 2.6 cm wide scintillator strips, read-out by means of WLS fibres. The fibres from the 6.7 m long strips were collected into four groups of 64 and coupled to 64-channel Hamamatsu H7546 photomultipliers. The muon spectrometers were made by an iron-core dipole magnet with drift tubes used as precision trackers and 22 layers of resistive plate chambers (RPC) [22] inside the magnetised iron. The 2D read-out was performed by means of 2.6 cm pitch and 8 m long vertical strips, which measured the coordinate in the bending plane, and 3.5 cm pitch and 8.7 m long horizontal strips, measuring the orthogonal coordinate.

The analysis presented here is based on TT and RPC data recorded during about five years from January 2008 to March 2013. In the TT, the trigger condition required either hits in the horizontal and vertical views of at least two planes or at least 10 hits in a single plane with a signal greater than 30 photo-electrons for 2008 and 2009 runs; from 2010 up to 2013 the latter requirements were lowered to 4 hits and 10 photo-electrons, respectively. Data from the RPCs of each spectrometer were acquired in presence of at least 3 planes fired in a time window of 200 ns. Events were recorded in presence of at least 5 TT and/or RPC hits in each view, horizontal and vertical, within a time window of 500 ns.

The TT systems were operative during most of 2008–2013, while the RPCs had a lower run-time, being operative only during CNGS runs and switched off during the CNGS winter shutdowns. Further details about the electronic detectors used for the cosmic ray muon analysis can be found in [18].

2 Cosmic ray muon flux measurement and its modulation

Cosmic ray induced events in the OPERA detector were selected outside of the CNGS spill time window. Once an event was tagged as “off-beam” it was classified as cosmic and processed in a dedicated way. The standard reconstruction package “OpRec” [23] was complemented with a set of algorithms developed for the different cosmic and beam event topologies. The reconstruction was effective at identifying single and multiple muon tracks (muon bundles). For this analysis, a total of about 4 million single muon events were selected requiring a single track reconstructed in both views.

Different atmospheric muon rates were measured in periods with and without RPC acquisition. The scale factor between the two rates was evaluated directly from data over the full data taking period, extracting the two constant terms $I_{\text{TT+RPC}}^0 = (3359 \pm 5) \mu/\text{day}$ and $I_{\text{TT-only}}^0 = (1960 \pm 5) \mu/\text{day}$ from a maximum likelihood fit on the two data sets.

In figure 1 (top panel) the flux of atmospheric single muons measured from 1 January 2008 (day 1 in the plot) to March 2013 is shown. After data quality cuts, the data set is composed of 1274 live days, out of which 919 days with TT+RPC. The longer downtime period corresponds to the first 5 months of 2009, when the acquisition was first stopped due to a DAQ upgrade and then to the earthquake in L’Aquila. Other shorter downtime periods are present in winter due to maintenance operations. Data with TT-only acquisition were rescaled to the TT+RPC average rate. The flux was fitted to

$$I_\mu(t) = I_\mu^0 + I_\mu^1 \cos \frac{2\pi}{T}(t - \phi) \quad (2.1)$$

The presence of a sinusoidal component with a period $T = (359 \pm 2)$ days is observed, with

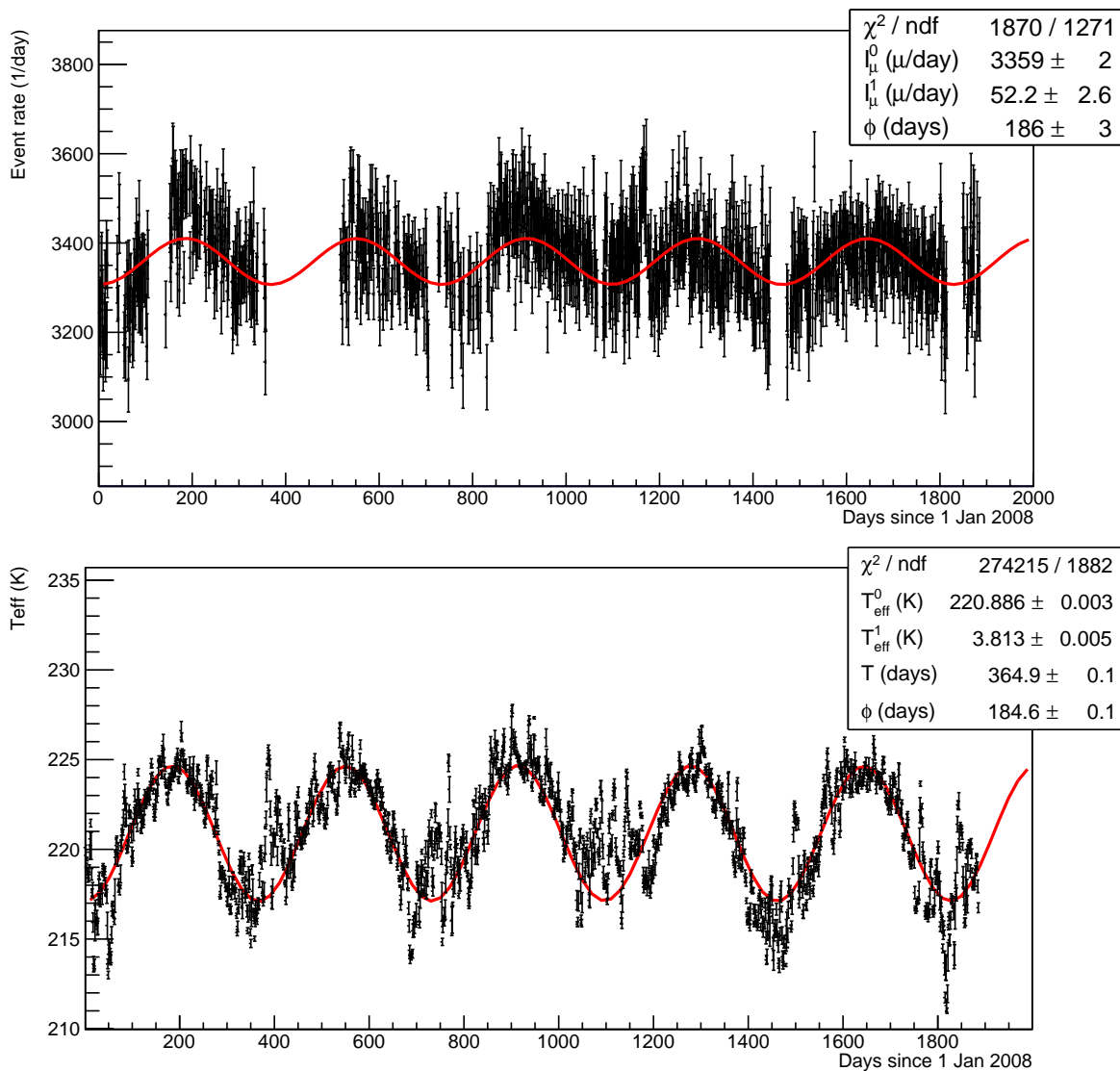


Figure 1. Single muon rate measured by the OPERA detector (top) and effective atmospheric temperature (bottom) from January 2008 to March 2013. Fit results are superimposed to the data sets. Symbols in the legends are defined in the text and in eqs. (2.1) and (3.1).

an amplitude amounting to $\alpha = I_\mu^1/I_\mu^0 = (1.55 \pm 0.08)\%$ of the average flux, and a phase $\phi = (197 \pm 5)$ days, with a $\chi^2/\text{dof} = 1.46$. In accordance with other LNGS experiments and with the fit result on the temperature reported in section 3, we take as the best estimate of the phase the result obtained fixing the period to one year of 365 days. The maximum is then observed at day $\phi = (186 \pm 3)$, corresponding to July 5, with a χ^2/dof value of 1.47. The result is also shown in figure 1 (top panel).

The Lomb-Scargle periodogram [24, 25] is a common tool for analysing the spectral content of unevenly spaced data to detect a periodic variation independently of the modulation phase. For the analysis presented here, we have exploited the generalised Lomb-Scargle periodogram, proposed in [26], which takes into account the non-zero average value of the event rate. The periodogram obtained for the single muon event rate is shown in figure 2

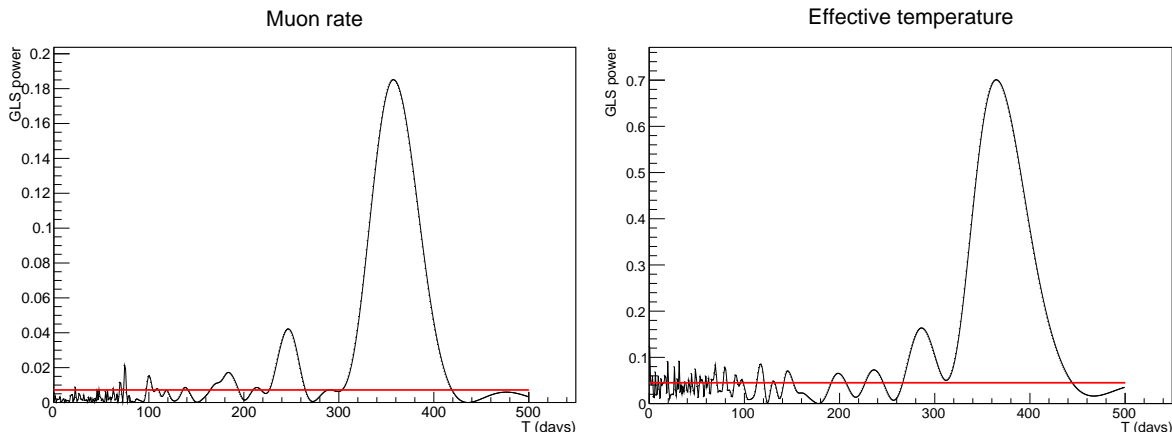


Figure 2. Generalised Lomb-Scargle periodograms for the measured muon rate (left) and the effective atmospheric temperature (right). The 99% significance level is also drawn as a reference in both periodograms.

(left panel). To assess the significance of the periodogram peaks, 10^5 toy experiments with a constant rate of $3359 \mu/\text{day}$ in the detector have been simulated, and the corresponding periodograms reconstructed. In figure 2 the 99% significance level, defined as the value for which 99% of the toy experiments result in a lower spectral power, is also shown.

The most significant peak is around one year (P_{max} at $T \sim 365$ days), but other less significant peaks are also present, as a consequence of the fact that eq. (2.1) is an approximation. A simulated experiment was performed extracting daily rates according to the result of our fit and comparing the periodograms obtained with and without data in the days of detector downtime. It shows that the amplitude of the peaks around 200 days increases with the detector downtime. A similar effect is also observed for the measured atmospheric temperature, treated in the next section, removing the data corresponding to days of detector downtime.

We checked possible systematic effects on the phase ϕ coming from the data taking stability along the years. Possible variations in DAQ efficiency are likely to coincide with calendar years by the fact that RPC were turned on and off on a yearly basis and DAQ maintenance and interventions were done during the CNGS winter shutdowns. By applying scale factors on a yearly basis, a constant average rate is achieved by definition. The normalisation, different at the few per mille level, was applied to TT-only and TT+RPC data, which were then rescaled one to the other. The rate was fitted to eq. (2.1) fixing the period to 365 days, the modulation maximum results on day $\phi = (177 \pm 3)$ with a $\chi^2/\text{dof} = 1.64$. Comparing these results with those obtained with a constant scale factor, we evaluate a systematic error on the phase as the semi-difference between the two ϕ values extracted with period fixed to one year, i.e. $\delta\phi_{\text{sys}} = 5$ days. Our best estimate of the muon rate maximum, obtained at fixed period $T = 365$ days, is found on day $\phi = 186 \pm 3_{\text{stat}} \pm 5_{\text{sys}}$, i.e. July 5.

3 Atmospheric temperature modulation

To measure the atmospheric temperature modulation, we used data from the European Centre for Medium-range Weather Forecasts (ECMWF) [27]. The Centre provides temperature values at different altitudes above given locations, obtained by means of data assimilation: a numerical procedure for combining a variety of different observations uniformly and consis-

tently within a global atmospheric model. The coordinates are those used in [7]: 13.5333° E, 42.4275° N. The atmospheric temperature is provided at 37 discrete pressure levels ranging from 1 to 1000 hPa four times in each day (0.00 h, 6.00 h, 12.00 h and 18.00 h). Averaging these temperature values using weights accounting for the production of pions and kaons at different altitudes (see appendix A), the effective atmospheric temperature T_{eff} has been calculated four times a day. The four measurements are then averaged and the variance used as an estimate of the uncertainty on the mean value.

The weights used in this analysis were computed as described in [28] and in previous experimental papers ([2–15]). The weights depend on: the inclusive meson production in the forward region; on the attenuation lengths of cosmic ray primaries, pions and kaons; as well as on the average value $\langle E_{\text{thr}} \cos \theta \rangle$, where E_{thr} is the minimum energy required for a muon to reach the considered underground site and θ is the angle between the muon and the vertical direction. From a full Monte Carlo simulation taking into account the rock map above Hall C of LNGS [29], it follows that $E_{\text{thr}} = 1.4 \text{ TeV}$ and $\langle E_{\text{thr}} \cos \theta \rangle = (1.1 \pm 0.2) \text{ TeV}$. A previous analysis [7] quoted a higher value, $E_{\text{thr}} = 1.8 \text{ TeV}$, extracted from numerical methods assuming a flat overburden [28]. All other parameters in the weight functions are site independent. Further details about the effective atmospheric temperature calculation are given in appendix A.

The T_{eff} values are shown in figure 1 (bottom panel) from January 2008 to March 2013. The temperature was fitted to a sinusoidal function similar to eq. (2.1):

$$T_{\text{eff}}(t) = T_{\text{eff}}^0 + T_{\text{eff}}^1 \cos \frac{2\pi}{T}(t - \phi) \quad (3.1)$$

The fit results are also shown in figure 1. The average effective temperature is 220.9 K, and a modulation is observed with an amplitude $\alpha = T_{\text{eff}}^1/T_{\text{eff}}^0 = (1.726 \pm 0.002)\%$ of T_{eff}^0 . The period $T = (364.9 \pm 0.1)$ days and phase $\phi = (184.6 \pm 0.1)$ days are similar to those observed for the single muon rate. A more refined study about the time correlation between temperature and muon rate is presented in the next section. In figure 2 (right panel) the generalised Lomb-Scargle periodogram is displayed also for T_{eff} . As for the muon rate, the most significant peak is around 365 days and other less significant peaks are present.

4 Cosmic ray flux and effective atmospheric temperature correlation

The possible presence of a time shift τ between the modulated components of the cosmic ray muon rate and of the effective atmospheric temperature has been investigated using the cross correlation function defined as:

$$R(\tau) = \int_0^{\Delta t} \frac{I_\mu(t) - I_\mu^0}{I_\mu^1} \frac{T_{\text{eff}}(t - \tau) - T_{\text{eff}}^0}{T_{\text{eff}}^1} \frac{dt}{\Delta t} \simeq \frac{1}{N_d} \sum_i \frac{I_\mu(t_i) - I_\mu^0}{I_\mu^1} \frac{T_{\text{eff}}(t_i - \tau) - T_{\text{eff}}^0}{T_{\text{eff}}^1} \quad (4.1)$$

I_μ^0 and T_{eff}^0 are the average values as obtained in the fits of section 2, for the single muon cosmic events, and of section 3, for the effective atmospheric temperature, while I_μ^1 and T_{eff}^1 are the corresponding amplitudes of the modulated components. The sum runs over the N_d days with both measurements.

The experimental cross correlation function is shown in black in figure 3. The two considered time series are correlated with no significant time delay. A peak with maximum at $\tau = 0$ days is observed, and we investigated with a toy Monte Carlo study if the peak is entirely due to the seasonal variation common to the two time series. In the toy Monte Carlo

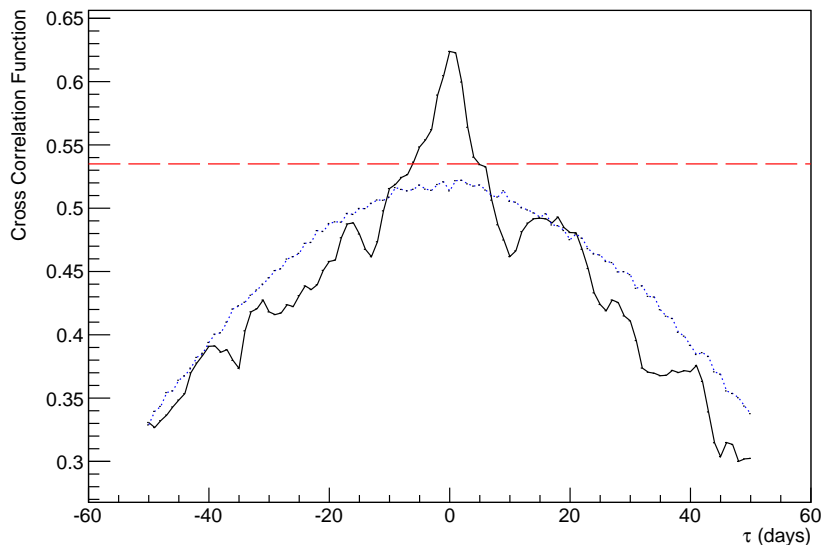


Figure 3. Cross correlation function (continuous black) between the measured daily muon rate and the effective atmospheric temperature. In dotted blue the result of one of the toy Monte Carlo experiments described in the text is reported, where the muon rate and the effective temperature have been extracted according to the fit results, but with equal time period and phase. In dashed red the 99% significance level for a pure seasonal correlation is also shown (see text for the definition).

simulation the muon rate and temperature data series are both modelled by pure sinusoidal variations with parameters $T = 365$ days and $\phi = 185$ days, according to the results of the fits reported in figure 1. We produced 10^5 Monte Carlo time series pairs; the cross correlation function of one of these pairs is shown in figure 3 by the dotted blue line, as a reference. The dashed red line represents the 99% C.L. for the pure seasonal cross correlation, defined as the value for which 99% of the toy cross correlation functions, evaluated at $\tau = 0$ days, fluctuate below it. The experimental peak exceeds this expected level and we conclude that the two time series appear to be cross correlated and synchronous both on seasonal and on short time scales.

In figure 4 the percentage deviation of the single muon flux, $\Delta I_\mu / I_\mu^0 = (I_\mu - I_\mu^0) / I_\mu^0$, is shown as a function of the relative effective temperature variation, $\Delta T_{\text{eff}} / T_{\text{eff}}^0 = (T_{\text{eff}} - T_{\text{eff}}^0) / T_{\text{eff}}^0$. According to the model described in appendix A, a proportionality relation is expected:

$$\frac{\Delta I_\mu(t)}{I_\mu^0} = \alpha_T \frac{\Delta T_{\text{eff}}(t)}{T_{\text{eff}}^0} \quad (4.2)$$

The effective temperature coefficient α_T depends on the energy threshold for atmospheric muons to reach the detector and on the ratio of kaon/pion production in cosmic rays interactions with the atmosphere. The linear fit performed on the plot of figure 4 gives $\alpha_T = 0.95 \pm 0.04$, consistent with expectations and with results from other LNGS experiments ([2–9]). In the fit, a constant term, a_0 , has been added to verify the consistency of our data with eq. (4.2). Its measured value $a_0 = -0.08 \pm 0.05$ is consistent with zero, as expected. The correlation coefficient is $R = 0.50$. It is worth noticing that R is proportional to the cross correlation function evaluated for $\tau = 0$, and can be obtained from eq. (4.1) replacing I_μ^1 and T_{eff}^1 , the modulation amplitudes, by the respective standard deviations.

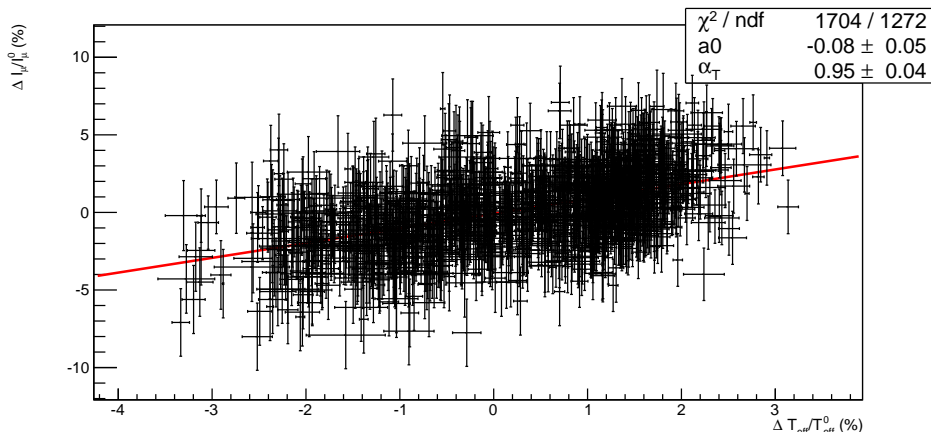


Figure 4. Correlation between the muon rate and the effective temperature relative variations. Symbols in the legend are defined in the text.

Two sources of systematic errors have been investigated: the energy threshold on cosmic ray muons collected by the detector, affecting the calculation of the effective atmospheric temperatures, and the muon rate data-set rescaling. Using $\langle E_{\text{thr}} \cos \theta \rangle \sim 1.8 \text{ TeV}$, as done by other LNGS experiments, the weights have been re-evaluated resulting in an effective temperature which is, on average, 0.2 K higher, with no appreciable effect on α_T measurement and also on the other analyses reported here.

The evaluation of α_T was performed also applying scaling factors, different year by year, to the muon flux measurements, independently with and without the RPC system as done in section 2; possible effects due to small changes in detector settings and data acquisition would be taken into account. The obtained effective temperature coefficient $\alpha_T = (0.93 \pm 0.04)$ is compatible with the previously quoted value within statistical uncertainties. The systematic error on α_T was therefore neglected with respect to the statistical one.

In figure 5 our result is shown together with the values measured by other experiments as a function of the energy threshold $\langle E_{\text{thr}} \cos \theta \rangle$. The model accounting for pion and kaon contributions to the muon flux is represented by the continuous black line, where the kaon/pion ratio $r_{K/\pi} = Z_{NK}/Z_{N\pi} = 0.144$ is fixed to the value inferred by the muon charge ratio analysis [30].

5 Conclusions

In this paper we report on studies about the seasonal variation of the flux of single muons generated in cosmic rays interactions in the high atmosphere, as measured by the OPERA electronic detectors located in the LNGS underground laboratory. The observed dependence is approximated by a constant flux with a $(1.55 \pm 0.08)\%$ seasonal component modelled as a sinusoid with one year period and maximum at $(186 \pm 3_{\text{stat}} \pm 5_{\text{sys}})$ day, corresponding to July 5.

An effective atmospheric temperature has been defined as the average of the temperatures of the air column above the experimental site, weighted by the production probability of high energy muons, detectable by the underground detector. The effective atmospheric temperature shows a time variation similar to cosmic ray rate changes.

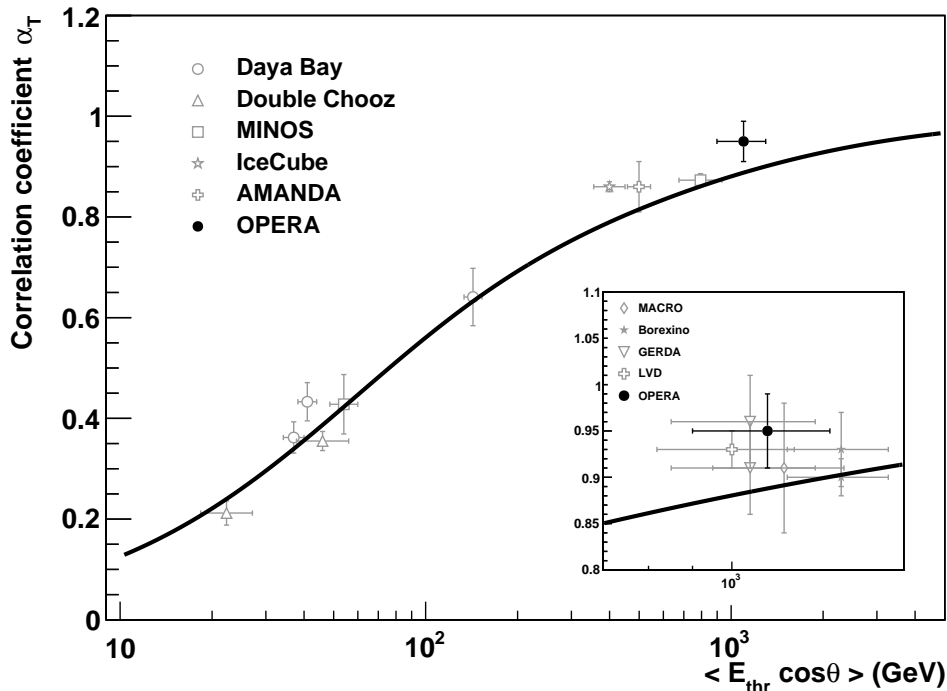


Figure 5. Experimental values of the correlation coefficient α_T as a function of $\langle E_{\text{thr}} \cos\theta \rangle$. The result of this analysis is displayed in black, measurements reported by other experiments in grey (AMANDA [10], IceCube [11], MINOS [12, 13], Double Chooz [14] and Daya Bay [15]). The seven measurements performed by five LNGS experiments (MACRO [3], Borexino [7, 8], GERDA [9], LVD [5, 6] and OPERA) are shown in the insert plot. The LNGS energy threshold — assumed to be (1.1 ± 0.2) TeV, as evaluated in this paper — is artificially displaced for each data point for the sake of visualisation. The continuous black line represents the model accounting for pions and kaons, with $r_{K/\pi} = 0.144$ [30].

A cross correlation function based study has demonstrated the presence of short term deviations from a seasonal cycle model, both in the cosmic ray flux and in the effective temperature. These short term deviations and the modulated components appear to co-vary between the two time series.

The effective atmospheric temperature and muon rate variations are positively correlated ($R = 0.50$). The effective temperature coefficient is measured to be $\alpha_T = (0.95 \pm 0.04)$, consistent with the model of π and K production in primary cosmic ray interactions and with previous experimental results. Both the phase and the effective temperature coefficient were measured by several LNGS experiments over different time periods in different decades, and the results agree very well within the quoted uncertainties.

A Models for modulation of cosmic rays and effective atmospheric temperature

In this appendix the correlation between the annual modulations of the temperature and of the cosmic ray single muon flux is discussed according to models described in [28].

Isolated muons are detected by underground experiments as a consequence of decays of (π^\pm and K^\pm) mesons produced by interactions of cosmic ray primaries in the atmosphere. An increase of the temperature results in a decrease of the air density and in a higher mean free

path of (π^\pm and K^\pm) mesons, with an increase of single muon flux, related to the temperature variation ΔT , given by:

$$\Delta I_\mu(t) = \int_0^\infty dX W(X) \Delta T(X, t) \quad (\text{A.1})$$

where the integral extends over the atmospheric depth and $W(X)$ reflects the altitude dependence of the mesons production and their decays into muons that can be observed underground.

A suitable model [28] for the purpose of this paper describes the atmosphere as an isothermal body with an effective temperature T_{eff} defined as:

$$T_{\text{eff}}(t) = \frac{\int_0^\infty dX T(X, t) W(X)}{\int_0^\infty dX W(X)} \simeq \frac{\sum_{n=0}^N \Delta X_n T(X_n, t) W(X_n)}{\sum_{n=0}^N \Delta X_n W(X_n)} \quad (\text{A.2})$$

where the approximation reflects the fact that the temperature is reported at discrete atmospheric levels X_n (not equally spaced).

Eq. (A.1) can be re-written as:

$$\Delta I_\mu(t) = \Delta T_{\text{eff}}(t) \int_0^\infty dX W(X) \quad (\text{A.3})$$

or

$$\frac{\Delta I_\mu(t)}{I_\mu^0} = \alpha_T \frac{\Delta T_{\text{eff}}(t)}{T_{\text{eff}}^0} \quad (\text{A.4})$$

with α_T , the effective temperature coefficient, defined as

$$\alpha_T = \frac{T_{\text{eff}}^0}{I_\mu^0} \int_0^\infty dX W(X) \quad (\text{A.5})$$

T_{eff}^0 being the effective temperature value for which $I_\mu(t) = I_\mu^0$.

The weight $W(X)$ used in eq. (A.1) can be expressed as the sum $W_\pi + W_K$, with W_π and W_K being the contributions of pions and kaons to the muon flux:

$$W^{\pi, K}(X) \simeq \frac{(1 - X/\Lambda'_{\pi, K})^2 e^{-X/\Lambda_{\pi, K}} A_{\pi, K}^1}{\gamma + (\gamma + 1) B_{\pi, K}^1 K(X) (\langle E_{\text{thr}} \cos \theta \rangle / \epsilon_{\pi, K})^2} \quad (\text{A.6})$$

where:

$$K(X) = \frac{(1 - X/\Lambda'_{\pi, K})^2}{(1 - e^{-X/\Lambda'_{\pi, K}}) \Lambda'_{\pi, K} / X} \quad (\text{A.7})$$

$A_{\pi, K}^1$ and $B_{\pi, K}^1$ are constants. A_π^1 is fixed to 1, while the other parameters depend on the masses of mesons and muons, as well as on the muon spectral index γ . A_K^1 is proportional to the ratio of forward muon production in decays of kaons and pions, $r_{K/\pi}$. $1/\Lambda'_{\pi, K} = 1/\Lambda_N - 1/\Lambda_{\pi, K}$, with Λ_N , Λ_π and Λ_K being the attenuation lengths for primaries, pions and kaons, respectively. E_{thr} is the energy required for a muon to reach the underground site and θ is the angle between the muon and the vertical direction. $\epsilon_{\pi, K}$ are the meson critical energies, for which decay and interaction have an equal probability.

The value of $\langle E_{\text{thr}} \cos \theta \rangle$ was estimated using a full Monte Carlo simulation taking into account the rock map above the Hall C of LNGS [29]. For OPERA $E_{\text{thr}} = 1.4 \text{ TeV}$ and $\langle E_{\text{thr}} \cos \theta \rangle = 1.1 \text{ TeV}$. All other parameters are site independent. The values used

Parameter	Value
A_π^1	1 [28]
A_K^1	0.38 $r_{K/\pi}$ [28]
$r_{K/\pi}$	0.149 [31]
γ	1.7 [31]
B_π^1	1.460 [28]
B_K^1	1.740 [28]
Λ_N	120 g/cm ² [31]
Λ_π	160 g/cm ² [31]
Λ_K	180 g/cm ² [31]
ϵ_π	0.115 TeV [31]
ϵ_K	0.850 TeV [31]

Table 1. Input parameters used for $W(X)$ evaluation.

in the present analysis are taken from [28, 31] and are reported in table 1 for the sake of completeness.

In figure 6 the five year (2008–2013) average temperature in the atmosphere for the LNGS site and the weight evaluated with $\langle E_{\text{thr}} \cos \theta \rangle = 1.1$ TeV as a function of the atmospheric pressure are shown. The weights are normalized to the highest value ($W(1 \text{ hPa}) = 0.0085748$). Higher atmosphere layers are given higher weights, since muons produced at low altitudes are on average less energetic and a larger fraction of them are below the threshold E_{thr} to cross the rock overburden.

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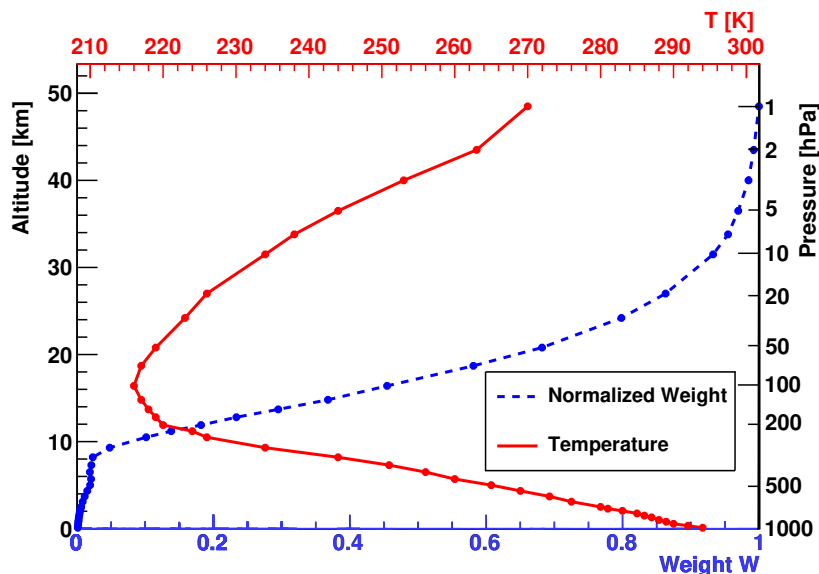


Figure 6. Average temperature (continuous red line) and normalized weight W (dashed blue line) as a function of pressure levels computed for the LNGS site (levels indicated by the dots). In the left vertical axis, the altitude corresponding to the pressure in the right vertical axis is shown; the conversion is extracted from [27] and the relationship is non-linear.

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