

The pulsation properties of λ bootis stars I. the southern TESS sample

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ABSTRACT

We analyse TESS light curves for 70 southern λ Boo stars to identify binaries and to determine which of them pulsate as δ Scuti stars. We find two heartbeat stars and two eclipsing binaries among the sample. We calculate that 81 per cent of λ Boo stars pulsate as δ Sct variables, which is about twice that of normal stars over the same parameter space. We determine the temperatures and luminosities of the λ Boo stars from photometry and *Gaia* DR2 parallaxes. A subset of 40 λ Boo stars have 2-min TESS data, reliable temperatures and luminosities, and δ Sct pulsation. We use Petersen diagrams (period ratios), échelle diagrams, and the period–luminosity relation to identify the fundamental mode in 20 of those 40 stars and conclude that a further 8 stars are not pulsating in this mode. For the remaining 12, the fundamental mode cannot be unambiguously identified. Further mode identification is possible for 12 of the fundamental mode pulsators that have regular sequences of pulsation overtones in their échelle diagrams. We use stellar evolution models to determine statistically that the λ Boo stars are only superficially metal weak. Simple pulsation models also better fit the observations at a metallicity of $Z = 0.01$ than at $Z = 0.001$. The TESS observations reveal the great potential of asteroseismology on λ Boo stars, for determining precise stellar ages and shedding light on the origin(s) of the λ Boo phenomenon.

Key words: asteroseismology – stars: chemically peculiar – stars: oscillations.

1 INTRODUCTION

The present era of space-based photometry has heralded a wave of advances in our understanding of stars. Thanks to asteroseismology, processes that are unconstrainable by other means, such as angular momentum transport (Aerts, Mathis & Rogers 2019), the extent of overshooting around the convective core (Johnston et al. 2019), and the existence of strong magnetic fields in stellar cores (Stello et al. 2016), are all now much better understood. Many of these results were made possible by the *Kepler* Mission (Borucki et al. 2010; Gilliland et al. 2010), which was designed to stare at a single, fixed patch of sky. The Transiting Exoplanet Survey Satellite (TESS; Ricker et al. 2015) differs, in that it surveys the sky in segments of 27-d duration, that will cover 85 per cent of the sky during the 2-yr nominal mission. This allows different problems to be addressed by targeting particular classes of stars. TESS is already enabling research on a broad range of stellar astrophysics, including compact pulsators (Bell et al. 2019, Charpinet et al. 2019),

chemically peculiar variable stars (Cunha et al. 2019; Sikora et al. 2019), massive stars (Pedersen et al. 2019) and δ Sct stars (Antoci et al. 2019; Bedding et al. 2020). In this paper, we use TESS to investigate λ Boo stars.

The λ Boo stars have spectral types from late-B to early-F and their absorption spectra carry the fingerprints of selective accretion (Venn & Lambert 1990). They are heavily depleted in refractory (iron-peak) elements by up to 2 dex (Stürenburg 1993), but have solar abundances of the volatile elements C, N, O, and S (Paunzen et al. 1999a; Kamp et al. 2001). Venn & Lambert (1990) hypothesized that the peculiarities originate when circumstellar dust grains are separated from gas, and only the latter is accreted by the star. However, the origin of the accreted material is a long-standing puzzle with several contenders that include residual circumstellar or protoplanetary discs (Kama, Folsom & Pinilla 2015), overdense pockets of the interstellar medium (Kamp & Paunzen 2002; Martínez-Galarza et al. 2009), and ablated hot Jupiters (Jura 2015). Since λ Boo stars appear to span a wide range of stellar ages, there might be multiple processes responsible for the phenomenon (Murphy & Paunzen 2017). Many young λ Boo stars have been observed to have protoplanetary discs (Malfait et al. 1998; Fisher et al. 2000; Meeus et al. 2001; Mawet et al.

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2017; Ligi et al. 2018; Cugno et al. 2019; Gratton et al. 2019; Macías et al. 2019; Bruzzone et al. 2020; Pérez et al. 2020), though the fraction of λ Boo stars with debris discs is not statistically different from that of normal stars (Gray et al. 2017). Disc accretion, potentially moderated by embedded planets, gives photospheric abundances in good agreement with observations (Jermyn & Kama 2018).

An important piece of the λ Boo puzzle is the question of whether λ Boo stars are only superficially metal weak (that is, they are peculiar stars), rather than being globally metal weak like the Population II stars. Here, asteroseismology is a valuable tool, probing the stellar interior rather than just the surface. The λ Boo stars are found in and around the instability strips of δ Sct and γ Dor variables, which pulsate in pressure and gravity modes, respectively (Aerts, Christensen-Dalsgaard & Kurtz 2010), and some also show Rossby modes (Saio et al. 2018a). For the γ Dor stars, *Kepler* has been transformational, with 611 class members identified and their pulsation properties characterized (Van Reeth et al. 2015; Li et al. 2020). However, the shorter TESS data sets do not have the frequency resolution required for the same level of study. Hence, in this work, we use the TESS data to identify and characterize the intersection of the λ Boo stars and δ Sct stars.

The δ Sct instability strip extends from mid-A to early-F spectral types (Dupret et al. 2004; Murphy et al. 2019), but the pulsational driving and damping mechanisms are not fully understood. Is not possible to determine a priori which of the theoretically possible oscillation modes will be excited, for a given stellar temperature, luminosity, and metallicity. In addition, there are many δ Sct stars located beyond the theoretical blue edge of the instability strip (Bowman & Kurtz 2018). A particular curiosity is that even in the middle of the δ Sct instability strip, only 70 percent of A stars pulsate as δ Sct stars, and this number declines rapidly towards the instability strip edges (Murphy et al. 2019). For the λ Boo stars, the pulsator fraction is expected and observed to be higher, because the He II partial ionization zone that drives the δ Sct pulsations is enhanced in helium from the accreted gas (Paunzen et al. 2002; Murphy 2014).

The fundamental mode in normal δ Sct stars follows a period–luminosity (P–L) relation (Poleski et al. 2010; McNamara 2011; Ziaali et al. 2019; Jayasinghe et al. 2020), but it is not strict. A colour term is sometimes included to account for the relatively wide temperature range of δ Sct stars which, at least in theory, excites different oscillation modes (Dupret et al. 2004). Paunzen et al. (2002) showed that λ Boo stars follow the same period–luminosity–colour (PLC) relation as normal δ Sct stars, and a separate metallicity term was not needed. By computing pulsation Q values, they also showed that λ Boo stars generally pulsate in high-overtone modes. Bedding et al. (2020) have found that λ Boo stars are overrepresented among their sample of δ Sct stars pulsating with regular overtone sequences at high frequencies.

TESS data for southern λ Boo stars constitute the most homogeneous, precise, and continuous time-series photometry available for the class. We describe these data in Section 2, along with the stellar parameters that allow us to place these stars in the HR diagram. These data offer the first opportunity to systematically assess the pulsation properties of λ Boo stars. We analyse the pulsator fraction and compare it to that of normal A stars in Section 3. We analyse the Fourier transforms of λ Boo light curves and attempt mode identification in Section 4. In Section 5, we readdress the question of whether the λ Boo stars are globally metal poor. A summary is given in Section 6 and notes on the pulsation properties of individual stars are given in Appendix A.

2 DATA COLLECTION AND PROCESSING

2.1 Target selection

The 70 λ Boo stars studied here originate from four sources: 27 confirmed λ Boo stars from the Murphy et al. (2015) catalogue, omitting the probable or uncertain members, and omitting HD 198161 which is not resolved from its neighbour HD 198160 in TESS pixels; 20 λ Boo stars from the Tycho sample in Gray et al. (2017); nine from the ‘Misc’ sample in the same paper; and 14 from a survey for southern λ Boo stars from Siding Spring Observatory (Murphy et al., submitted). A few targets belong to more than one source.

2.2 TESS light curves

TESS observes each sector for 27 d and the sectors have overlap near the ecliptic poles. Stars can therefore fall in more than one sector and have light curves of longer durations (in some multiple of 27 d). Of our 70 targets, 21 have observations in more than one sector. The majority (56) of our TESS targets were observed at 2-min cadence and the rest only have full-frame image (FFI) data at 30-min cadence. We used all available TESS sectors for each λ Boo star unless they had 2-min data in some sectors but only FFI data in others, in which case we analysed only the 2-min data.

Jenkins et al. (2016) described the TESS science processing operations centre (SPOC) pipeline that produces the 2-min light curves. We downloaded these 2-min data using the PYTHON package LIGHTKURVE (Lightkurve Collaboration 2018). We used ‘PDCSAP’ flux, which generally produces smooth light curves without any large drifts in flux from systematic effects on the detector. We performed no additional processing on these light curves.

For the 14 λ Boo stars that only have 30-min FFI data, SPOC light curves are not currently available for download. We used the PYTHON package ELEANOR (Feinstein et al. 2019) to extract time-series from the FFIs, and kept data with *quality flag* = 0. We made light curves using the ‘corrected’ flux, which we found to produce smoother light curves than either the ‘raw’ or ‘psf’ flux options. There were often drifts in flux following the data downlink events, typically lasting for ~ 1 d. We excluded these data points from the analysis if we deemed them severe because their inclusion generates strong peaks at low frequencies in the Fourier transforms of the light curves, which can mask lower-amplitude pulsations.

Using the ASTROPY.STATS package, we computed the Lomb-Scargle periodogram of each TESS light curve, from which we extracted the frequencies and amplitudes of the strongest three peaks above 4 d^{-1} (Table 1). We considered peaks with a signal-to-noise ratio (SNR) above 4 to be significant and did not tabulate peaks below this threshold. We computed the SNR by dividing the amplitude of the peak by the average amplitude in a 2-d^{-1} window around that peak. Of the 56 λ Boo stars with 2-min TESS data, 47 pulsate as δ Sct stars and only 40 of those have reliable temperatures and luminosities (Section 2.3). We show the Fourier transforms of those 40 stars in Figs 1 and 2, ordered by luminosity. Fourier transforms for all 70 stars are provided in Appendix B with the extracted peaks in Table 1 highlighted.

Peaks at frequencies below 4 d^{-1} are not pressure modes and are not included in Table 1. However, if such peaks were present, we reported them in the notes on individual stars (Appendix A). Although substantial progress has recently been made on the interpretation of g and r modes in A/F stars (Kurtz et al. 2015; Van Reeth, Tkachenko & Aerts 2016; Saio et al. 2018a,b; Li et al. 2019; Ouazzani et al. 2019), this requires light curves with a time-

Table 1. Pulsation properties, including frequencies (f_i) and amplitudes (a_i), of the 70 λ Boo stars in this study. The final three columns indicate the 14 stars with only 30-min FFI data, the upper limit on any undetected pulsation modes for the 16 λ Boo stars that are not δ Sct pulsators, and the TESS sectors in which each target was observed. Note that HD 198160 and HD 198161 are separated by less than one TESS pixel and are considered as a single object in this paper (see also Appendix A).

Name	f_1 d ⁻¹	a_1 mmag	f_2 d ⁻¹	a_2 mmag	f_3 d ⁻¹	a_3 mmag	FFI only	limit mmag	TESS Sectors
HD 319	–	–	–	–	–	–	–	0.01	2
HD 3922	8.542	6.487	5.020	5.158	8.720	4.754	–	–	1–2
HD 4158	8.715	3.328	7.106	2.578	9.169	1.233	–	–	3
HD 6870	17.873	1.873	15.429	1.490	15.860	1.252	–	–	1–2
HD 7908	–	–	–	–	–	–	–	0.02	3
HD 11413	26.804	8.450	25.215	8.342	25.107	3.303	–	–	2–3
HD 13755	7.925	7.172	12.662	4.989	15.080	4.850	–	–	2–3
HD 17341	24.451	2.419	23.276	2.247	20.768	1.654	–	–	3–4
HD 23392	–	–	–	–	–	–	–	0.02	4
HD 28490	5.232	7.406	6.948	4.068	7.487	2.836	–	–	5
HD 28548	61.517	3.573	65.246	3.475	57.559	2.341	–	–	5
HD 30422	48.091	1.243	51.242	0.922	41.531	0.839	–	–	5
HD 31295	41.390	0.634	44.999	0.184	37.196	0.177	–	–	5
HD 31508	19.040	1.447	20.278	0.973	14.321	0.861	–	–	1–13
HD 36726	–	–	–	–	–	–	–	0.04	6
HD 37411	–	–	–	–	–	–	FFI	0.2	6
HD 38043	14.004	2.800	14.510	1.348	15.034	0.854	–	–	1–3, 5–6, 8–13
HD 41958	12.429	8.161	6.390	7.815	11.754	7.478	–	–	1, 4–5, 7–8, 11
HD 42503	7.178	4.315	5.250	4.231	6.185	3.219	–	–	5–7
HD 43533	23.558	0.267	30.937	0.076	26.193	0.043	–	–	6–7
HD 44930	18.152	2.924	23.994	2.436	25.418	1.067	–	–	6
HD 46722	20.764	8.981	43.360	5.084	20.233	4.377	–	–	6–7
HD 47425	14.201	8.945	15.841	8.648	5.931	3.292	–	–	6–7
HD 68695	17.172	2.078	13.827	0.438	17.172	0.167	FFI	–	7–8
HD 73211	17.588	0.089	21.346	0.059	5.324	0.056	–	–	8
HD 74423	8.757	1.901	–	–	–	–	–	–	8–11
HD 75654	14.340	3.289	15.240	0.401	13.017	0.216	–	–	8–9
HD 76097	8.933	11.666	9.872	4.747	8.800	3.237	–	–	8
HD 80426	11.016	10.076	39.305	0.622	31.271	0.431	–	–	8
HD 83041	14.528	1.060	14.085	0.314	15.862	0.295	–	–	8–9
HD 84159	9.139	43.124	10.900	8.558	8.003	7.580	–	–	8
HD 88554	5.648	0.081	7.939	0.070	–	–	FFI	–	9–10
HD 94326	10.130	0.778	17.607	0.428	10.383	0.404	–	–	9–10
HD 94390	–	–	–	–	–	–	–	0.02	9–10
HD 98069	13.094	3.529	15.426	2.615	16.225	2.067	FFI	–	9
HD 100546	–	–	–	–	–	–	FFI	0.2	11
HD 101412	–	–	–	–	–	–	FFI	0.08	10–11
HD 102541	19.898	7.326	19.490	1.126	37.793	1.060	–	–	10
HD 103701	6.188	3.504	4.653	3.399	12.376	0.695	–	–	10
HD 109738	33.769	3.623	12.083	1.688	33.328	1.443	–	–	11
HD 111786	15.023	1.728	31.020	1.690	36.006	1.622	–	–	10
HD 112682	–	–	–	–	–	–	–	0.04	10
HD 112948	29.755	3.449	31.825	2.332	27.040	2.113	–	–	10
HD 119896	–	–	–	–	–	–	FFI	0.05	11
HD 120122	8.555	0.113	16.066	0.093	–	–	–	–	11
HD 125508	16.066	1.028	9.257	0.949	8.931	0.736	–	–	11
HD 126627	–	–	–	–	–	–	–	0.04	11–12
HD 127659	6.537	4.144	6.884	3.351	11.904	2.442	–	–	11–12
HD 139614	34.777	1.557	21.365	1.142	28.241	0.881	–	–	12
HD 153747	21.471	6.266	19.879	2.862	22.809	2.283	–	–	12
HD 154153	–	–	–	–	–	–	–	0.015	12
HD 154951	9.335	0.228	–	–	–	–	FFI	–	12
HD 159021	13.825	0.931	10.707	0.819	14.156	0.521	–	–	12–13
HD 162193	24.068	7.926	19.342	3.286	20.566	1.311	–	–	13
HD 168740	30.757	4.565	12.576	3.528	24.082	2.639	–	–	13
HD 168947	14.594	2.339	17.644	1.518	15.242	0.929	FFI	–	13
HD 169142	–	–	–	–	–	–	FFI	0.15	13
HD 169346	23.441	1.341	20.120	0.994	21.225	0.905	FFI	–	13
HD 184779	12.566	9.401	13.837	4.812	22.875	3.315	–	–	13
HD 188164	–	–	–	–	–	–	FFI	0.015	13

Table 1 – continued

Name	f_1 d ⁻¹	a_1 mmag	f_2 d ⁻¹	a_2 mmag	f_3 d ⁻¹	a_3 mmag	FFI only	limit mmag	TESS Sectors
HD 198160	35.961	0.967	33.302	0.945	41.137	0.767	–	–	13
HD 203709	14.110	2.362	16.083	1.803	6.277	1.787	–	–	1
HD 210111	30.207	1.211	18.593	1.192	16.233	1.128	–	–	1
HD 213669	28.242	12.647	14.017	10.228	15.430	7.520	–	–	1
HD 214582	19.763	0.542	18.185	0.214	–	–	–	–	1
HD 223352	–	–	–	–	–	–	–	0.02	2
HD 290799	22.534	4.955	38.940	2.059	21.734	0.947	–	–	6
HD 294253	–	–	–	–	–	–	–	0.05	6
BD-11 1239	13.150	1.630	15.835	0.844	10.189	0.765	FFI	–	6
T Ori	10.961	0.496	–	–	–	–	FFI	–	6

base of at least a few years for enough frequency resolution to study the pulsation modes. We considered λ Boo stars that have data from Year 1 of TESS observations (Sectors 1–13). Figs 1 and 2 show power excesses at these frequencies in many stars, but their detailed analysis is not possible given the relatively short durations of the light curves.

2.3 Stellar parameters

The position of λ Boo stars on the HR diagram is important in many contexts. If λ Boo stars are young and accreting material left over from star or planet formation then they should lie near the ZAMS. In addition, and regardless of their age, their position with respect to the δ Sct instability strip is important for understanding mode excitation and for an asteroseismic interpretation of the global stellar metallicity and age. Finally, accurate luminosities allow us to use the period–luminosity relation to identify the fundamental mode in some stars.

To determine stellar luminosities, we used *Gaia* DR2 parallaxes where available (Lindgren et al. 2018). Two program stars have no parallax data in either of the *Gaia* and *Hipparcos* catalogues. We still discuss the pulsation properties of these stars, but they are not included in analyses where a luminosity is required (Sections 3–5).

We determined T_{eff} from Geneva 7-colour photometry for the 44 stars for which this was available, since this approach results in the smallest uncertainties. To calibrate the Geneva photometry, we used the grids published by Kunzli et al. (1997), which are valid for B- to G-type stars. Note that for stars cooler than B9, no reddening-free indices exist within this photometric system. Therefore, to get reliable temperature values, reddening values must be provided as input. We also determined T_{eff} from spectral energy distributions (SEDs) using the VOSA (Virtual Observatory Sed Analyser) tool v6.0 (Bayo et al. 2008), which was also the primary source of T_{eff} for targets without Geneva photometry. SEDs offer a precision of ~ 250 K. No outliers were detected among the 44 targets with T_{eff} determined by both methods: the two methods agree to within 1σ for 67 per cent of the overlapping sample. Our final adopted uncertainties on the derived effective temperatures were fixed to be at least 2 per cent, which is typical of the uncertainty set by the accuracy of measured interferometric angular diameters (Casagrande et al. 2014; White et al. 2018).

The extinction in the V band, A_V , was computed using the same photometry that we used to calculate the effective temperatures. For the 27 stars where this photometry was lacking, we used the python package MWDUST (Bovy et al. 2016) to extract extinctions

from the combined (default) dust map using the target coordinates and distances.

Using the stellar temperatures, parallaxes, A_V , and V -band bolometric corrections, we calculated bolometric luminosities using the ‘direct mode’ of the ISOCCLASSIFY package (Huber et al. 2017). We adopted 0.03 mag uncertainties in reddening and bolometric corrections. To estimate stellar masses, we fitted the effective temperatures and luminosities derived in the previous step to MIST isochrones (MESA Isochrones and Stellar Tracks; Choi et al. 2016; Dotter 2016) using the ‘grid mode’ of *isoclassify*, assuming a solar-neighbourhood metallicity prior. The procedure also yielded estimates of surface gravities, which were used for the interpolation of bolometric corrections in the previous step. To ensure consistency we performed one iteration between the ‘direct mode’ and ‘grid mode’ calculations. The resulting stellar parameters are given in Table 2.

We identified four binary systems in the sample from their TESS light curves: HD 31508 and HD 94390, which we discovered to be heartbeat stars (Appendix A); and HD 168947 and HD 74423, which we found to be eclipsing binaries (the latter was also discussed by Handler et al. 2020). We searched the Washington Double Star (WDS) catalogue (Mason et al. 2001) for close binaries (<10 -arcmin separation), and found three objects with magnitude differences less than 2.5 mag (HD 15164, HD 198160, and HD 170680). Finally, we also used the renormalized unit weight error (RUWE) from *Gaia* DR2 to identify probable binaries. Objects with $\text{RUWE} > 2.0$ are likely to be binaries (Evans 2018; Rizzuto et al. 2018) and three such binaries were identified, all of which were identified in our WDS search. Other quality controls, such as a filter on targets with more than 1.0 mag of extinction, did not identify any additional stars whose parameters might be questionable. We plot the HR diagram of 58 stars in Fig. 3, excluding from the original 70 stars the 7 stars found to be binaries, four additional stars above 10 000 K that are too hot to be δ Sct pulsators (HD 23392, HD 100546, HD 127659, and HD 294253), and HD 169142, which has anomalous stellar parameters. It is this sample of 58 stars that we use in Sections 3–5; the boolean column ‘good’ in Table 2 indicates which stars passed the filter.

3 PULSATOR FRACTION ANALYSIS

It has been known for some time that the fraction of λ Boo stars showing δ Sct pulsations is high, but this has been difficult to quantify. Collecting precise, continuous, and lengthy time-series photometry from the ground is difficult, requiring multisite coordination and being limited by the effects of Earth’s atmosphere.

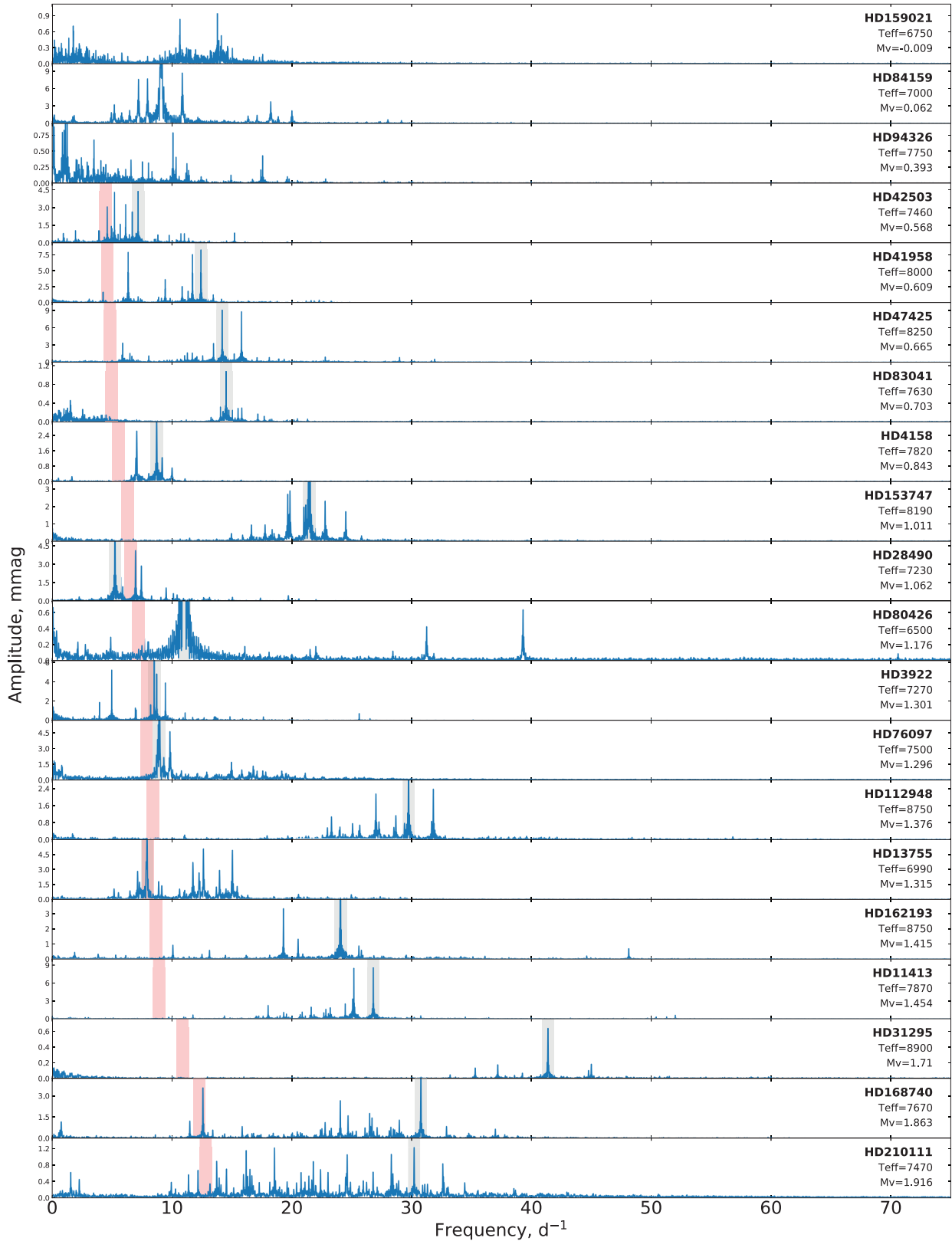


Figure 1. Fourier amplitude spectra of TESS 2-min data for 20 λ Boo stars that are also δ Sct stars, ordered by luminosity (brightest at the top). Only stars with well-determined atmospheric parameters are shown (see Section 2.3). The strongest peak (Table 1) is indicated with a grey rectangle; the red rectangles give the predicted location of the fundamental mode based on the period–luminosity relation for δ Sct stars (Section 4). Each panel has its own vertical axis range, which aids visibility of smaller peaks, occasionally at the expense of taller peaks. Asterisks denote stars studied in Bedding et al. (2020) and double asterisks denote stars with échelle diagrams shown in that paper. The Fourier transforms of all stars are plotted individually in Appendix B. Continued in Fig. 2.

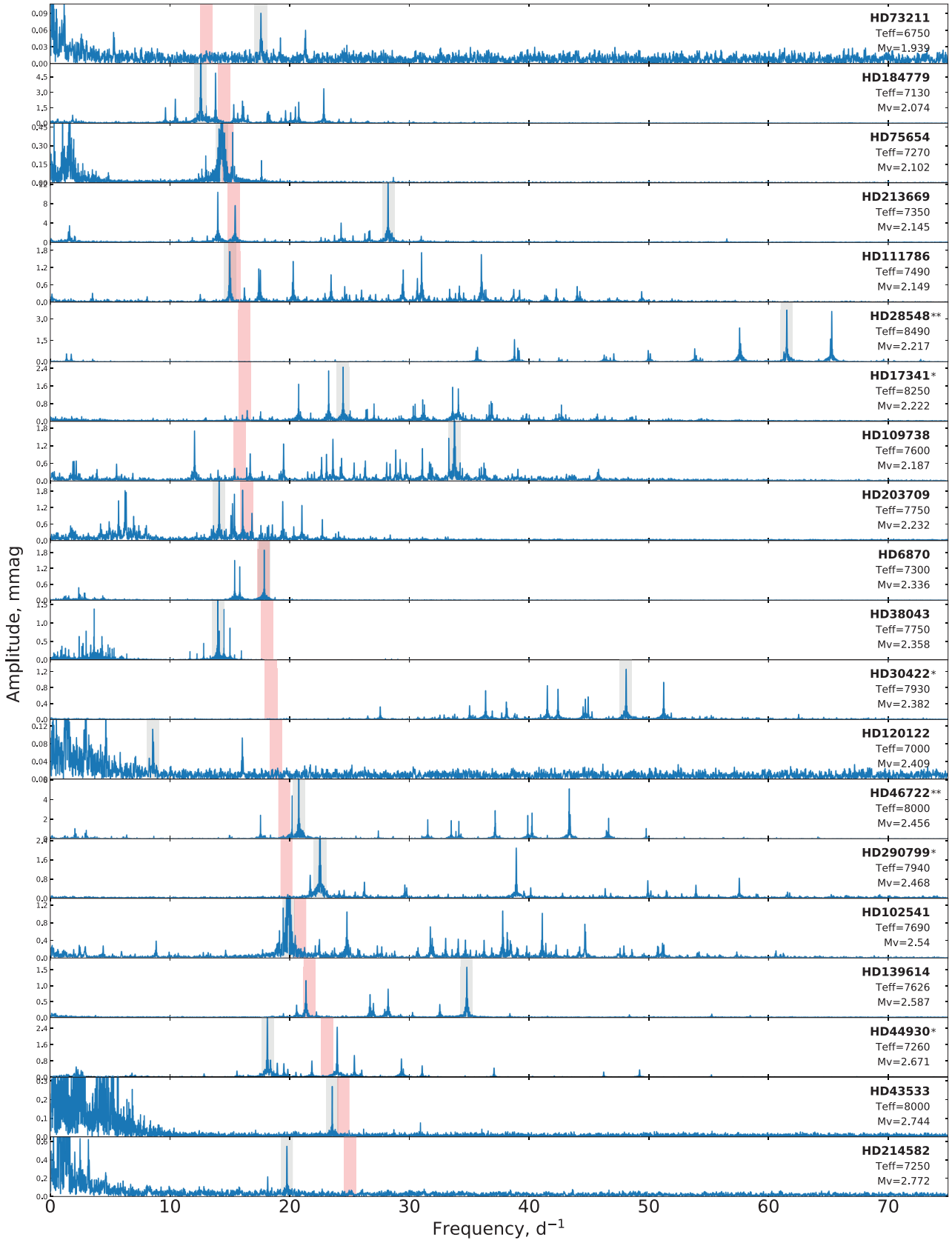


Figure 2. Same as Fig. 1, for the less luminous half of the sample. The Fourier transforms of all stars are plotted individually in Appendix B.

Table 2. Stellar parameters for the southern λ Boo stars used in this paper. Columns are: the HD number or other identifier, V -band magnitude, effective temperature, luminosity, absolute V magnitude, surface gravity, metallicity, mass, extinction in the V band, Renormalized Unit Weight Error, a Boolean flag indicating whether the parameters are ‘good’ (see the text), the magnitude difference from any companion in the WDS catalogue, and the corresponding WDS separation.

HD	V mag	T_{eff} K	L L_{\odot}	M_V mag	$\log g$	[Fe/H]	M M_{\odot}	A_V mag	RUWE	good	Δm mag	Sep. arcsec
HD 319	5.92	8080 \pm 162	23.72 ^{+1.13} _{-1.06}	1.28 ^{+0.05} _{-0.05}	3.95 ^{+0.04} _{-0.04}	0.00 ^{+0.12} _{-0.10}	2.02 ^{+0.06} _{-0.05}	0.01	1.33	1	5.10	2.20
HD 3922	8.34	7270 \pm 145	22.57 ^{+1.02} _{-0.96}	1.30 ^{+0.05} _{-0.05}	3.77 ^{+0.04} _{-0.03}	-0.01 ^{+0.08} _{-0.09}	1.96 ^{+0.04} _{-0.04}	0	1.01	1	-	-
HD 4158	9.54	7820 \pm 156	34.12 ^{+2.06} _{-1.93}	0.84 ^{+0.06} _{-0.06}	3.77 ^{+0.04} _{-0.04}	0.02 ^{+0.07} _{-0.07}	2.17 ^{+0.04} _{-0.04}	0.58	0.98	1	-	-
HD 6870	7.45	7300 \pm 146	8.69 ^{+0.36} _{-0.35}	2.34 ^{+0.04} _{-0.04}	4.11 ^{+0.04} _{-0.05}	-0.01 ^{+0.15} _{-0.16}	1.61 ^{+0.08} _{-0.09}	0	1.08	1	-	-
HD 7908	7.30	7120 \pm 142	5.94 ^{+0.25} _{-0.23}	2.73 ^{+0.04} _{-0.04}	4.20 ^{+0.04} _{-0.05}	-0.02 ^{+0.12} _{-0.16}	1.48 ^{+0.07} _{-0.08}	0	1.02	1	-	-
HD 11413	5.93	7870 \pm 157	19.31 ^{+0.80} _{-0.76}	1.45 ^{+0.04} _{-0.04}	3.98 ^{+0.04} _{-0.05}	0.02 ^{+0.12} _{-0.22}	1.95 ^{+0.06} _{-0.13}	0.01	0.93	1	-	-
HD 13755	7.83	6990 \pm 140	21.96 ^{+0.94} _{-0.89}	1.31 ^{+0.04} _{-0.05}	3.73 ^{+0.04} _{-0.04}	0.09 ^{+0.09} _{-0.21}	1.99 ^{+0.04} _{-0.16}	0	0.99	1	-	-
HD 17341	9.32	8250 \pm 250	9.90 ^{+0.48} _{-0.45}	2.22 ^{+0.05} _{-0.05}	4.25 ^{+0.04} _{-0.05}	-0.14 ^{+0.13} _{-0.16}	1.66 ^{+0.07} _{-0.08}	0.10	1.01	1	-	-
HD 23392	8.24	10140 \pm 203	28.69 ^{+1.61} _{-1.52}	1.37 ^{+0.06} _{-0.06}	4.26 ^{+0.03} _{-0.04}	-0.24 ^{+0.10} _{-0.15}	2.14 ^{+0.05} _{-0.09}	0.16	1.24	0	-	-
HD 28490	9.53	7230 \pm 145	27.96 ^{+1.91} _{-1.77}	1.06 ^{+0.07} _{-0.07}	3.68 ^{+0.05} _{-0.04}	-0.09 ^{+0.10} _{-0.08}	1.98 ^{+0.04} _{-0.04}	0.05	0.89	1	-	-
HD 28548	9.22	8490 \pm 170	10.04 ^{+0.48} _{-0.46}	2.22 ^{+0.05} _{-0.05}	4.29 ^{+0.03} _{-0.03}	-0.25 ^{+0.13} _{-0.15}	1.64 ^{+0.07} _{-0.07}	0.12	1.06	1	-	-
HD 30422	6.17	7930 \pm 159	8.42 ^{+0.35} _{-0.33}	2.38 ^{+0.04} _{-0.04}	4.25 ^{+0.04} _{-0.04}	-0.13 ^{+0.12} _{-0.14}	1.60 ^{+0.06} _{-0.07}	0	0.90	1	-	-
HD 31295	4.65	8900 \pm 231	16.91 ^{+1.42} _{-1.29}	1.71 ^{+0.09} _{-0.09}	4.22 ^{+0.04} _{-0.05}	-0.12 ^{+0.14} _{-0.16}	1.90 ^{+0.08} _{-0.09}	0.09	0	1	-	-
HD 31508	9.57	6750 \pm 250	28.88 ^{+1.54} _{-1.45}	1.04 ^{+0.06} _{-0.06}	3.60 ^{+0.04} _{-0.05}	0.03 ^{+0.11} _{-0.12}	2.05 ^{+0.06} _{-0.14}	0.05	1.07	0	-	-
HD 36726	8.85	9950 \pm 418	40.73 ^{+2.47} _{-2.31}	0.95 ^{+0.06} _{-0.06}	4.13 ^{+0.06} _{-0.10}	-0.07 ^{+0.15} _{-0.18}	2.36 ^{+0.09} _{-0.16}	0.29	0.92	1	-	-
HD 37411	9.79	11462 \pm 229	-	-	-	-	-	1.56	2.92	0	-	-
HD 38043	9.45	7750 \pm 250	8.53 ^{+0.41} _{-0.39}	2.36 ^{+0.05} _{-0.05}	4.21 ^{+0.05} _{-0.06}	-0.07 ^{+0.13} _{-0.15}	1.61 ^{+0.07} _{-0.08}	0.06	1.08	1	-	-
HD 41958	8.78	8000 \pm 250	43.23 ^{+2.11} _{-1.98}	0.61 ^{+0.05} _{-0.05}	3.69 ^{+0.06} _{-0.05}	-0.14 ^{+0.13} _{-0.10}	2.19 ^{+0.06} _{-0.05}	0.21	1.11	1	-	-
HD 42503	7.43	7460 \pm 149	44.53 ^{+1.98} _{-1.87}	0.57 ^{+0.05} _{-0.05}	3.59 ^{+0.03} _{-0.03}	0.00 ^{+0.07} _{-0.07}	2.27 ^{+0.04} _{-0.03}	0	1.04	1	-	-
HD 43533	9.61	8000 \pm 250	6.05 ^{+0.28} _{-0.27}	2.74 ^{+0.05} _{-0.05}	4.31 ^{+0.04} _{-0.04}	-0.28 ^{+0.14} _{-0.19}	1.44 ^{+0.06} _{-0.06}	0.14	1.13	1	-	-
HD 44930	9.42	7260 \pm 145	6.41 ^{+0.35} _{-0.33}	2.67 ^{+0.06} _{-0.06}	4.21 ^{+0.04} _{-0.05}	-0.04 ^{+0.12} _{-0.15}	1.50 ^{+0.07} _{-0.08}	0.06	0.72	1	-	-
HD 46722	9.29	8000 \pm 250	7.86 ^{+0.35} _{-0.33}	2.46 ^{+0.05} _{-0.05}	4.27 ^{+0.04} _{-0.05}	-0.16 ^{+0.12} _{-0.14}	1.57 ^{+0.05} _{-0.07}	0.07	0.96	1	-	-
HD 47425	9.53	8250 \pm 250	41.06 ^{+2.57} _{-2.40}	0.67 ^{+0.07} _{-0.07}	3.78 ^{+0.11} _{-0.07}	-0.14 ^{+0.35} _{-0.14}	2.16 ^{+0.25} _{-0.06}	0.19	0.85	1	-	-
HD 68695	9.87	8436 \pm 169	57.47 ^{+2.21} _{-2.12}	-	3.76 ^{+0.03} _{-0.04}	0.16 ^{+0.07} _{-0.08}	2.55 ^{+0.04} _{-0.04}	0.10	1.08	1	-	-
HD 73211	9.02	6750 \pm 250	12.73 ^{+0.63} _{-0.59}	1.94 ^{+0.05} _{-0.05}	3.84 ^{+0.07} _{-0.06}	0.01 ^{+0.15} _{-0.15}	1.71 ^{+0.09} _{-0.08}	0.64	1.16	1	-	-
HD 74423	8.61	8250 \pm 250	69.30 ^{+3.74} _{-3.54}	0.10 ^{+0.06} _{-0.06}	3.65 ^{+0.04} _{-0.05}	0.11 ^{+0.09} _{-0.09}	2.62 ^{+0.05} _{-0.05}	0.10	1.12	0	-	-
HD 75654	6.37	7270 \pm 145	11.00 ^{+0.45} _{-0.43}	2.10 ^{+0.04} _{-0.04}	4.03 ^{+0.05} _{-0.05}	0.00 ^{+0.15} _{-0.17}	1.68 ^{+0.09} _{-0.09}	0	0.92	1	-	-
HD 76097	9.56	7500 \pm 250	22.56 ^{+1.38} _{-1.28}	1.30 ^{+0.06} _{-0.06}	3.83 ^{+0.08} _{-0.06}	-0.03 ^{+0.19} _{-0.14}	1.94 ^{+0.12} _{-0.06}	0.35	1.06	1	-	-
HD 80426	9.16	6500 \pm 250	25.75 ^{+1.50} _{-1.41}	1.18 ^{+0.06} _{-0.06}	3.60 ^{+0.05} _{-0.07}	0.13 ^{+0.11} _{-0.13}	2.07 ^{+0.06} _{-0.11}	0.89	1.39	1	-	-
HD 83041	8.79	7630 \pm 930	39.03 ^{+1.91} _{-1.79}	0.70 ^{+0.05} _{-0.05}	3.71 ^{+0.23} _{-0.11}	0.00 ^{+0.15} _{-0.16}	2.23 ^{+0.10} _{-0.11}	0.66	1.07	1	-	-
HD 84159	9.51	7000 \pm 250	70.08 ^{+5.68} _{-5.20}	0.06 ^{+0.08} _{-0.09}	3.36 ^{+0.06} _{-0.05}	-0.04 ^{+0.17} _{-0.17}	2.29 ^{+0.08} _{-0.09}	0.48	1.05	1	-	-
HD 88554	9.32	6920 \pm 153	30.63 ^{+1.63} _{-1.53}	0.96 ^{+0.06} _{-0.06}	3.60 ^{+0.03} _{-0.03}	0.02 ^{+0.10} _{-0.09}	2.06 ^{+0.05} _{-0.05}	0.14	1.24	1	-	-
HD 94326	7.76	7750 \pm 155	52.25 ^{+2.41} _{-2.27}	0.39 ^{+0.05} _{-0.05}	3.60 ^{+0.03} _{-0.02}	-0.03 ^{+0.11} _{-0.09}	2.28 ^{+0.05} _{-0.05}	0.05	0.93	1	-	-
HD 94390	8.96	7000 \pm 250	21.11 ^{+1.00} _{-0.94}	1.37 ^{+0.05} _{-0.05}	3.75 ^{+0.06} _{-0.07}	0.04 ^{+0.14} _{-0.18}	1.94 ^{+0.07} _{-0.14}	0.42	1.03	0	-	-
HD 98069	8.16	7500 \pm 250	37.8 ^{+2.09} _{-1.98}	0.73 ^{+0.06} _{-0.06}	3.68 ^{+0.05} _{-0.05}	0.06 ^{+0.10} _{-0.10}	2.24 ^{+0.05} _{-0.06}	0.06	1.03	1	-	-
HD 100546	6.30	10720 \pm 214	33.71 ^{+1.44} _{-1.36}	1.31 ^{+0.04} _{-0.04}	4.29 ^{+0.03} _{-0.03}	-0.37 ^{+0.10} _{-0.12}	2.17 ^{+0.04} _{-0.05}	0.20	1.19	0	10.10	4.50
HD 101412	9.29	8101 \pm 162	25.43 ^{+1.27} _{-1.19}	1.17 ^{+0.05} _{-0.05}	3.93 ^{+0.04} _{-0.04}	0.00 ^{+0.11} _{-0.09}	2.04 ^{+0.06} _{-0.05}	0.09	0.95	1	-	-
HD 102541	7.94	7690 \pm 154	7.24 ^{+0.32} _{-0.30}	2.54 ^{+0.05} _{-0.05}	4.25 ^{+0.03} _{-0.04}	-0.12 ^{+0.11} _{-0.14}	1.54 ^{+0.06} _{-0.07}	0.09	1.11	1	-	-
HD 103701	8.73	7000 \pm 250	41.41 ^{+3.39} _{-3.09}	0.66 ^{+0.08} _{-0.09}	3.56 ^{+0.05} _{-0.06}	0.13 ^{+0.11} _{-0.13}	2.31 ^{+0.06} _{-0.15}	0.29	2.04	0	-	-
HD 109738	8.30	7600 \pm 152	9.83 ^{+0.43} _{-0.41}	2.19 ^{+0.05} _{-0.05}	4.14 ^{+0.04} _{-0.04}	-0.02 ^{+0.14} _{-0.16}	1.67 ^{+0.08} _{-0.09}	0.09	0.90	1	-	-
HD 111786	6.14	7490 \pm 150	10.26 ^{+0.45} _{-0.42}	2.15 ^{+0.05} _{-0.05}	4.11 ^{+0.05} _{-0.04}	-0.01 ^{+0.15} _{-0.16}	1.68 ^{+0.09} _{-0.09}	0.02	1.98	1	-	-
HD 112682	9.67	8000 \pm 250	45.62 ^{+3.63} _{-3.32}	0.52 ^{+0.08} _{-0.08}	3.68 ^{+0.08} _{-0.06}	-0.14 ^{+0.16} _{-0.11}	2.19 ^{+0.08} _{-0.05}	0.26	1.04	1	-	-
HD 112948	9.35	8750 \pm 250	22.52 ^{+2.19} _{-1.95}	1.38 ^{+0.10} _{-0.10}	4.11 ^{+0.06} _{-0.06}	-0.04 ^{+0.14} _{-0.16}	2.02 ^{+0.09} _{-0.10}	0.07	1.13	1	-	-
HD 119896	8.22	6580 \pm 132	20.15 ^{+1.23} _{-1.15}	1.45 ^{+0.06} _{-0.06}	3.65 ^{+0.04} _{-0.04}	0.03 ^{+0.10} _{-0.08}	1.89 ^{+0.05} _{-0.05}	0.37	1.00	1	-	-

Table 2 – continued

HD	V mag	T_{eff} K	L L_{\odot}	M_V mag	$\log g$	[Fe/H]	M M_{\odot}	A_V mag	RUWE	good	Δm mag	Sep. arcsec
HD 120122	9.11	7000 \pm 250	8.18 $^{+0.44}_{-0.42}$	2.41 $^{+0.06}_{-0.06}$	4.06 $^{+0.07}_{-0.08}$	0.00 $^{+0.15}_{-0.17}$	1.57 $^{+0.09}_{-0.09}$	0.07	1.15	1	–	–
HD 125508	9.14	7250 \pm 250	–	–	–	–	–	–	16.63	0	–	–
HD 126627	9.00	7070 \pm 141	5.91 $^{+0.26}_{-0.24}$	2.75 $^{+0.05}_{-0.05}$	4.19 $^{+0.04}_{-0.05}$	– 0.01 $^{+0.13}_{-0.16}$	1.48 $^{+0.07}_{-0.08}$	0.09	0.90	1	–	–
HD 127659	9.31	11900 \pm 435	109.42 $^{+5.14}_{-4.84}$	0.28 $^{+0.05}_{-0.05}$	4.05 $^{+0.05}_{-0.05}$	– 0.23 $^{+0.12}_{-0.14}$	2.85 $^{+0.07}_{-0.06}$	1.34	0.93	0	–	–
HD 139614	8.24	7626 \pm 153	6.78 $^{+0.30}_{-0.28}$	2.59 $^{+0.05}_{-0.05}$	4.26 $^{+0.03}_{-0.04}$	– 0.12 $^{+0.11}_{-0.14}$	1.52 $^{+0.06}_{-0.08}$	0.02	1.09	1	3.53	1.40
HD 153747	7.42	8190 \pm 164	30.11 $^{+1.43}_{-1.34}$	1.01 $^{+0.05}_{-0.05}$	3.88 $^{+0.05}_{-0.04}$	– 0.06 $^{+0.12}_{-0.08}$	2.09 $^{+0.06}_{-0.04}$	0.12	0.90	1	–	–
HD 154153	6.18	7400 \pm 567	16.99 $^{+0.73}_{-0.69}$	1.60 $^{+0.04}_{-0.05}$	3.93 $^{+0.13}_{-0.15}$	– 0.01 $^{+0.15}_{-0.17}$	1.86 $^{+0.09}_{-0.10}$	0.34	1.30	1	–	–
HD 154951	8.78	7030 \pm 141	12.57 $^{+0.60}_{-0.57}$	1.94 $^{+0.05}_{-0.05}$	3.92 $^{+0.04}_{-0.05}$	0.00 $^{+0.15}_{-0.19}$	1.72 $^{+0.08}_{-0.10}$	0.23	0.98	1	–	–
HD 159021	8.67	6750 \pm 250	74.62 $^{+4.56}_{-4.27}$	– 0.01 $^{+0.06}_{-0.06}$	3.33 $^{+0.06}_{-0.05}$	– 0.02 $^{+0.20}_{-0.19}$	2.30 $^{+0.09}_{-0.10}$	0.05	1.09	1	–	–
HD 162193	8.66	8750 \pm 250	21.79 $^{+1.10}_{-1.03}$	1.42 $^{+0.05}_{-0.05}$	4.12 $^{+0.05}_{-0.05}$	– 0.04 $^{+0.14}_{-0.16}$	2.01 $^{+0.09}_{-0.10}$	0.09	1.14	1	–	–
HD 168740	6.12	7670 \pm 153	13.26 $^{+0.56}_{-0.53}$	1.86 $^{+0.04}_{-0.04}$	4.06 $^{+0.04}_{-0.05}$	0.00 $^{+0.15}_{-0.17}$	1.77 $^{+0.09}_{-0.09}$	0.02	0.98	1	–	–
HD 168947	8.11	7510 \pm 150	29.91 $^{+1.44}_{-1.35}$	1.00 $^{+0.05}_{-0.05}$	3.76 $^{+0.04}_{-0.06}$	0.14 $^{+0.10}_{-0.37}$	2.17 $^{+0.05}_{-0.24}$	0.12	0.95	0	–	–
HD 169142	8.16	6700 \pm 322	5.93 $^{+0.33}_{-0.31}$	2.87 $^{+0.14}_{-0.14}$	4.65 $^{+0.02}_{-0.03}$	– 0.79 $^{+0.13}_{-0.14}$	1.43 $^{+0.03}_{-0.03}$	0	0.86	0	6.50	0.10
HD 169346	9.27	7520 \pm 150	16.34 $^{+0.85}_{-0.80}$	1.66 $^{+0.05}_{-0.06}$	3.95 $^{+0.04}_{-0.04}$	– 0.01 $^{+0.17}_{-0.17}$	1.83 $^{+0.10}_{-0.10}$	0.13	0.82	1	–	–
HD 184779	8.90	7130 \pm 143	11.01 $^{+0.57}_{-0.53}$	2.07 $^{+0.05}_{-0.05}$	3.99 $^{+0.05}_{-0.04}$	0.00 $^{+0.15}_{-0.17}$	1.68 $^{+0.09}_{-0.09}$	0	1.02	1	–	–
HD 188164	6.37	8160 \pm 163	23.82 $^{+1.02}_{-0.97}$	1.26 $^{+0.04}_{-0.05}$	3.97 $^{+0.05}_{-0.04}$	– 0.02 $^{+0.17}_{-0.10}$	2.02 $^{+0.11}_{-0.05}$	0.08	1.18	1	–	–
HD 198160	6.21	7930 \pm 159	13.44 $^{+0.56}_{-0.53}$	1.85 $^{+0.04}_{-0.04}$	4.12 $^{+0.04}_{-0.05}$	– 0.02 $^{+0.14}_{-0.17}$	1.79 $^{+0.09}_{-0.09}$	0.02	1.11	0	0.35	3.20
HD 198161	6.55	7900 \pm 158	13.72 $^{+0.58}_{-0.55}$	1.85 $^{+0.04}_{-0.04}$	4.10 $^{+0.04}_{-0.05}$	– 0.01 $^{+0.15}_{-0.17}$	1.80 $^{+0.08}_{-0.10}$	0.02	1.24	0	–	–
HD 203709	9.58	7750 \pm 250	9.54 $^{+0.47}_{-0.45}$	2.23 $^{+0.05}_{-0.05}$	4.18 $^{+0.05}_{-0.06}$	– 0.05 $^{+0.13}_{-0.15}$	1.66 $^{+0.07}_{-0.09}$	0.07	0.98	1	–	–
HD 210111	6.39	7470 \pm 149	12.84 $^{+0.54}_{-0.51}$	1.92 $^{+0.04}_{-0.04}$	4.02 $^{+0.04}_{-0.04}$	– 0.01 $^{+0.16}_{-0.17}$	1.75 $^{+0.09}_{-0.09}$	0	0.98	1	–	–
HD 213669	7.42	7350 \pm 147	10.30 $^{+0.44}_{-0.42}$	2.15 $^{+0.04}_{-0.05}$	4.07 $^{+0.05}_{-0.05}$	0.00 $^{+0.15}_{-0.17}$	1.66 $^{+0.09}_{-0.08}$	0	1.10	1	–	–
HD 214582	9.48	7250 \pm 250	5.81 $^{+0.29}_{-0.27}$	2.77 $^{+0.05}_{-0.05}$	4.22 $^{+0.05}_{-0.06}$	– 0.07 $^{+0.14}_{-0.16}$	1.47 $^{+0.06}_{-0.08}$	0.02	1.21	1	–	–
HD 223352	4.57	9930 \pm 199	30.09 $^{+3.46}_{-3.00}$	1.27 $^{+0.11}_{-0.12}$	4.22 $^{+0.04}_{-0.04}$	– 0.14 $^{+0.11}_{-0.18}$	2.19 $^{+0.09}_{-0.13}$	0.06	0.00	1	7.03	3.40
HD 290799	10.80	7940 \pm 159	7.63 $^{+0.68}_{-0.62}$	2.47 $^{+0.09}_{-0.09}$	4.28 $^{+0.03}_{-0.04}$	– 0.18 $^{+0.12}_{-0.14}$	1.55 $^{+0.07}_{-0.08}$	0.02	1.29	1	–	–
HD 294253	9.67	10580 \pm 212	32.31 $^{+2.75}_{-2.50}$	1.33 $^{+0.09}_{-0.09}$	4.28 $^{+0.04}_{-0.03}$	– 0.33 $^{+0.11}_{-0.14}$	2.17 $^{+0.07}_{-0.08}$	0.23	1.18	0	–	–
BD-11 1239	9.70	7000 \pm 250	10.75 $^{+0.59}_{-0.56}$	2.11 $^{+0.06}_{-0.06}$	3.96 $^{+0.07}_{-0.08}$	0.00 $^{+0.15}_{-0.17}$	1.66 $^{+0.09}_{-0.09}$	0.15	1.12	1	–	–
T Ori	11.25	6980 \pm 140	9.06 $^{+1.22}_{-1.08}$	2.30 $^{+0.14}_{-0.14}$	4.03 $^{+0.07}_{-0.06}$	0.00 $^{+0.15}_{-0.17}$	1.58 $^{+0.10}_{-0.10}$	0.66	1.20	1	–	–

This resulted in heterogeneous surveys for variability, with short time-series of differing precision that did not include all group members. Despite these challenges, Paunzen et al. (2002) were able to conclude that the pulsator fraction of λ Boo stars is at least 70 per cent. It was difficult to assess whether this pulsator fraction is abnormally high due to the lack of a suitable reference against which to evaluate it because variability surveys of normal δ Sct stars suffered from the same issues, but in greater numbers. Now, thanks to *Kepler* and TESS, homogeneous samples exist that make meaningful comparisons possible.

To assess whether λ Boo stars have a higher pulsator fraction than normal A/F stars, we consider their position on the HR diagram using the stellar parameters determined in Section 2.3. Several λ Boo stars lie outside the observed and theoretical δ Sct instability strips, so it is unsurprising that they would not pulsate. Murphy et al. (2019) recently assessed the pulsator fraction of normal A/F stars in the *Kepler* field as a function of position on the HR diagram, providing a benchmark against which to assess the λ Boo stars.

The sensitivity of TESS is different from *Kepler* for two principal reasons: (i) the effective passband (governed by the CCD response function) is redder for TESS, which leads to TESS pulsation amplitudes being 16.5 per cent lower than their *Kepler* equivalents

(Lund 2019); and (ii) TESS light curves are shorter in duration so the Fourier noise is higher. To correct for this sensitivity difference, we estimated the average Fourier noise of *Kepler* and TESS light curves by first finding the median Fourier amplitude for each target (this is a good estimate of the noise level; Murphy et al. 2019), and then taking the median of the distribution of these noise levels. We found that the median noise level is 13 times greater in the TESS data. Following our significance criterion of SNR = 4 for detecting p modes in the TESS data, we concluded that *Kepler* pulsators would typically be detectable with TESS if the *Kepler* pulsation SNR is at least $(1.165 \times 13.3 \times 4) = 62.0$. Since the median *Kepler* noise is 0.00133 mmag, this corresponds to an amplitude limit of 0.082 mmag. That is, a peak with an observed amplitude above 0.082 mmag in *Kepler* data should be detectable in TESS data.¹ Adopting this amplitude limit, we found that 1826 of the

¹Note that this calculation does not correct for amplitude smearing caused by the integrations being a large fraction of the pulsation period in 30-min data (Chaplin et al. 2011; Murphy 2012); for our targets, which have a median f_1 value of 18 d^{-1} the correction is only 2.25 per cent. Further, for FFI data, the exposure times are almost identical to the *Kepler* ones so no correction would apply.

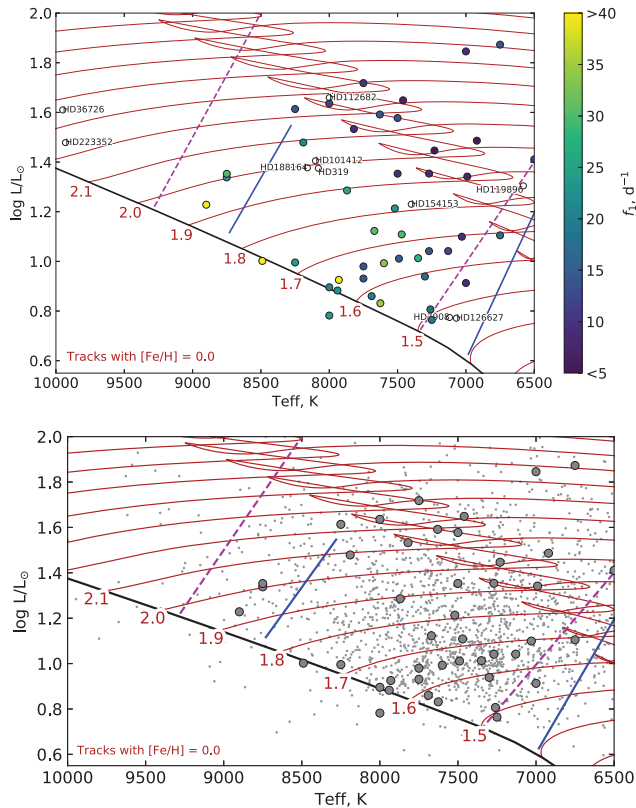


Figure 3. Top: HR diagram of the southern λ Boo stars observed by TESS. Empty circles show non-pulsators while pulsators are shown as filled circles colour-coded by the strongest oscillation frequency. Stars with $A_V > 1.0$ mag, RUWE > 2.0 , or a nearby bright companion in the WDS catalogue are not shown. Bottom: The same pulsating λ Boo stars are shown with the *Kepler* δ Sct stars for comparison, corrected for differences in sensitivity (see Section 3).

1988 *Kepler* δ Sct stars in the Murphy et al. (2019) sample would have been detectable at TESS precision, and we used this sample of 1826 *Kepler* stars as our comparison sample in the following pulsator fraction calculations.

Using the same method as Murphy et al. (2019), we used the 1826 δ Sct stars to calculate the fraction of pulsators as a function of position in the HR diagram. Using grid cells of approximately 200 K in T_{eff} and 0.1 dex in $\log L$, we computed the pulsator fraction at the position of each λ Boo star and treated this as the probability that the star would be observed to pulsate by TESS. Uncertainties were estimated by taking the standard deviation of the pulsator fraction in neighbouring cells, which are approximately 1σ away in the observed quantities, T_{eff} and $\log L$. Using this model, we computed an expected pulsator fraction of 0.41 ± 0.05 for our stars. The observed fraction is much higher: 47 of the 58 λ Boo stars with reliable atmospheric parameters are observed to pulsate, giving a pulsator fraction of 0.81. Thus the fraction of λ Boo stars that show δ Sct pulsations is about twice that of normal stars.

4 MODE IDENTIFICATION FOR λ BOO STARS

There has been significant recent progress in mode identification for δ Sct stars with the discovery of regular patterns of peaks in high-frequency δ Sct stars (Bedding et al. 2020). In these stars, it is possible to see ridges in an échelle diagram when the correct frequency separation, $\Delta\nu$, is used, and this separation is

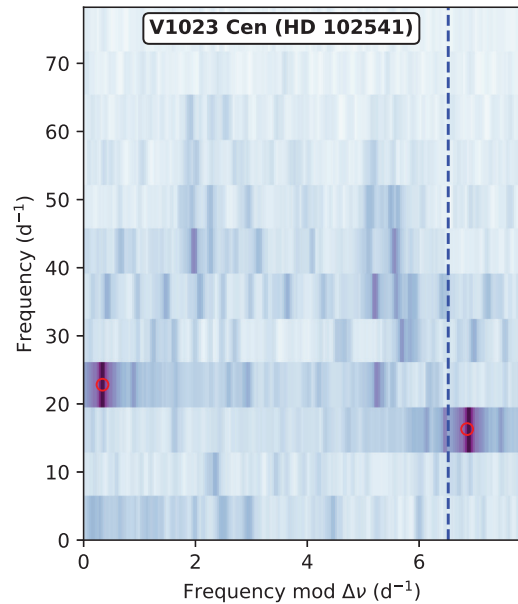


Figure 4. Échelle diagram of HD 102541, showing nearly vertical ridges at $x \approx 2$ and $x \approx 5.5$ d^{-1} . The region right of the dashed blue is duplicated for clarity. The red circle at the left of order 4 highlights the fundamental mode.

approximately proportional to the square root of the mean stellar density (Aerts et al. 2010). Bedding et al. (2020) have shown that ridges of $\ell = 0$ and $\ell = 1$ modes can be identified, and the radial order n can be inferred. They were able to match the pulsation modes to models to determine stellar ages much more precisely than by positioning the star on an HR diagram using its *Gaia* parallax.

Of particular interest is the fact that λ Boo stars are often high-frequency pulsators (Paunzen et al. 2002) and are also overrepresented in the Bedding et al. (2020) sample (6 out of 60). This makes application of the échelle diagram to the λ Boo stars in this paper attractive. Although we do not attempt to model individual oscillation modes here, identifying candidates for future asteroseismic modelling and establishing a provisional mode identification is an important first step.

While there remains progress to be made in interpreting the subtleties of the $\ell = 0$ and $\ell = 1$ ridges, the échelle diagrams can be used to identify the fundamental mode in some stars. For the 18 δ Sct stars in Bedding et al. (2020) for which the fundamental mode was identifiable, the frequency of this mode was close to three times the large separation (their extended data figure 2). In other words, the fundamental mode lies at the right of the third order or the left of the fourth order in the échelle. Twelve of the 40 stars in Figs 1 and 2 show ridges in their échelle diagrams that enabled $\Delta\nu$ to be measured and the fundamental mode to be determined. An example is shown in Fig. 4, created with the ECHELLE package (Hey & Ball 2020).

A second tool for identifying the fundamental mode in stars with well-determined parallaxes is via the period–luminosity relation, already discussed in the Introduction. In this work, we use a P–L relation to help identify the fundamental mode.

Finally, some δ Sct stars pulsate in low-overtone radial modes that have a known period ratio. Particularly for high-amplitude pulsators, period ratios have been used to identify the fundamental mode and the first- and second-overtone modes using so-called Petersen diagrams (Petersen & Christensen-Dalsgaard 1996). Broadly

speaking, the fundamental and first-overtone modes have a period ratio of ~ 0.77 and the first- and second-overtone modes have a ratio close to 0.80. However, the ratios depend on: (i) the stellar density, and by extension, on the period of the fundamental mode; (ii) the stellar metallicity, with metal-poor stars having slightly higher ratios (Petersen & Christensen-Dalsgaard 1996); (iii) the stellar rotation, which shifts the ratios by half a per cent even for modest rotation of $v = 50 \text{ km s}^{-1}$ (Suárez, Garrido & Goupil 2006); and (iv) near-degeneracy effects such as mode coupling, which imparts extra scatter in the ratios (Suárez, Garrido & Moya 2007).

When used alone, the above three methods sometimes leave considerable uncertainty in mode identification, but when two or more methods agree the identification is stronger. In addition, the fundamental mode usually has a high amplitude if it is excited, and it is the lowest frequency pressure mode that can be excited, though mixed modes and g modes can occur at lower frequencies. Hence, a dominant peak of low frequency ($\sim 4\text{--}25 \text{ d}^{-1}$) identifiable with one or more of the three methods above is likely to be the fundamental mode.

Using these tools, in Table 3 we provide the most likely identification for the fundamental mode in the 40 pulsating λ Boo stars shown in Figs 1 and 2. These are the stars with good parallaxes, on which the P–L relation is likely to be useful, and with 2-min data, so we can be sure that the observed frequencies are not Nyquist aliases. Not all stars pulsate in the fundamental mode, and in some stars there is no clear candidate from the many excited modes. Nonetheless, the fundamental mode was identifiable in 20 of the 40 stars, and for an additional 8 stars we can conclude that the fundamental mode is not excited. Our results show that the majority of λ Boo stars do pulsate in the fundamental mode.

5 ARE λ BOO STARS GLOBALLY METAL POOR?

Spectroscopy and multicolour photometry are sensitive to the metallicity of the photosphere, whereas evolutionary and pulsation models are sensitive to the global metallicity of the star. In combination, these techniques can determine whether λ Boo stars are superficially or globally metal weak. The most detailed analysis of the interior metallicity of a λ Boo star was done by Casas et al. (2009), who used asteroseismology to determine that 29 Cyg is globally metal poor. However, the models were quite poorly constrained, with the 1σ results spanning a mass range of 1.67 to $2.09 M_{\odot}$ in the solar-metallicity case, or 1.42 to $1.81 M_{\odot}$ in the globally metal-poor case. The Casas et al. (2009) result also remains at odds with the analysis by Paunzen et al. (2015), who computed stellar evolution and pulsation models of some λ Boo stars that had SuperWASP photometry, reaching the opposite conclusion. Since then, extremely precise distances have become available from *Gaia* DR2, allowing much better constrained models of λ Boo stars to be constructed.

We revisit the topic of global metallicities for λ Boo stars using a similar two-pronged methodology to Paunzen et al. (2015). Firstly, in Section 5.1 we compare the positions on the HR diagram of the 58 λ Boo stars having reliable atmospheric parameters (see Section 2.3) compared to stellar evolution models of two different metallicities. Secondly, in Section 5.2 we investigate pulsation models for a subsample of eight well-studied southern λ Boo stars (listed in Table 4), chosen for having detailed abundance analyses in the literature.

5.1 Stellar evolution models

We first consider the location of our target stars on the Hertzsprung–Russell (HR) diagram, with respect to evolutionary tracks of two

Table 3. Most likely identification of the fundamental mode for the 40 pulsating λ Boo stars in Figs 1 and 2. Stars are listed in the same order as the figures. The right column summarizes the method(s) used in the identification; more detail is given in the notes on individual stars in Appendix A. Some stars clearly do not pulsate in the fundamental mode ‘<not excited>’ and for some stars it is not possible to identify the mode ‘<unclear>’.

HD	Fundamental d^{-1}	Method(s)
HD 159021	10.707	Petersen
HD 84159	<unclear>	–
HD 94326	<unclear>	–
HD 42503	<unclear>	–
HD 41958	<unclear>	–
HD 47425	<unclear>	–
HD 83041	<unclear>	–
HD 4158	7.106	Petersen
HD 153747	<not excited>	P–L
HD 28490	6.948	échelle, P–L
HD 80426	<unclear>	–
HD 3922	<unclear>	–
HD 76097	<unclear>	–
HD 112948	<not excited>	P–L
HD 13755	7.925	P–L, Petersen
HD 162193	10.129	P–L
HD 11413	11.794	échelle
HD 31295	<not excited>	P–L
HD 168740	12.576	échelle, P–L
HD 210111	12.215	échelle, P–L
HD 73211	<not excited>	P–L, Petersen
HD 184779	12.568	P–L, Petersen
HD 75654	14.340	P–L
HD 213669	14.017	P–L
HD 111786	15.023	P–L
HD 28548	<not excited>	échelle, P–L
HD 17341	<unclear>	–
HD 109738	<unclear>	–
HD 203709	<unclear>	–
HD 6870	17.873	P–L
HD 38043	<not excited>	multiplets
HD 30422	<not excited>	échelle
HD 120122	<not excited>	P–L
HD 46722	20.764	échelle, P–L
HD 290799	22.534	échelle, P–L
HD 102541	19.898	échelle, P–L
HD 139614	21.365	échelle, P–L
HD 44930	18.152	échelle
HD 43533	23.558	P–L
HD 214582	19.763	P–L

different metallicities ($Z = 0.010$ and 0.001) and two values of overshooting from the hydrogen-burning convective core ($f_{\text{ov}} = 0.00$ and 0.02 ; Herwig 2000). We calculated the evolutionary tracks with MESA (Paxton et al. 2011, 2013, 2015, 2018, 2019) ver. 11554. We adopted an initial hydrogen abundance of $X_{\text{ini}} = 0.7$, and used OPAL opacity tables (Iglesias & Rogers 1996) supplemented with the data provided by Ferguson et al. (2005) for low temperatures. The solar chemical element mixture was used, as determined by Grevesse & Noels (1993). We used the MESA equation of state, which is a blend of the equation of state from OPAL (Rogers & Nayfonov 2002), SCVH (Saumon, Chabrier & van Horn 1995), PTEH (Pols et al. 1995), HELM (Timmes & Swesty 2000), and PC (Potekhin & Chabrier 2010). Nuclear reaction rates were from the JINA REACLIB data base (Cyburt et al. 2010) plus additional tabulated weak reaction rates (Fuller, Fowler & Newman 1985; Oda

Table 4. Parameters of the eight stars chosen for pulsational analysis. Columns are the Henry Draper catalogue number, visual magnitude, projected surface rotation velocity (references in final column), *Gaia* parallax (*Gaia* Collaboration 2016, 2018), effective temperature and luminosity (Table 2), stellar radius (calculated), and literature references to spectroscopic abundance analyses.

HD	V mag	$v \sin i$ km s^{-1}	π mas	$\log(T_{\text{eff}}/\text{K})$	$\log(L/L_{\odot})$	R R_{\odot}	Literature abundances
319	5.923	60	11.81 ± 0.14	3.9074 ± 0.0086	1.375 ± 0.020	2.48 ± 0.13	Stürenburg (1993), Paunzen et al. (1999a)
11413	5.929	125	12.726 ± 0.045	3.8960 ± 0.0086	1.286 ± 0.017	2.36 ± 0.12	Stürenburg (1993), Paunzen et al. (1999a)
31295	4.648	115	29.21 ± 0.53	3.9494 ± 0.0111	1.228 ± 0.033	1.73 ± 0.13	Venn & Lambert (1990), Stürenburg (1993) Paunzen et al. (1999a, 1999b), Kamp et al. (2001)
36726	8.84	80	3.017 ± 0.064	3.9978 ± 0.0179	1.610 ± 0.025	2.15 ± 0.20	Andrievsky et al. (2002)
75654	6.368	45	13.972 ± 0.032	3.8615 ± 0.0086	1.041 ± 0.017	2.09 ± 0.10	Solano et al. (2001), Kamp et al. (2001)
210111	6.386	30	12.714 ± 0.056	3.8733 ± 0.0086	1.109 ± 0.017	2.14 ± 0.10	Stürenburg (1993), Paunzen et al. (1999a) Solano et al. (2001)
290799	10.8	70	2.273 ± 0.063	3.8998 ± 0.0086	0.883 ± 0.036	1.46 ± 0.11	Andrievsky et al. (2002)
294253	9.67	70	2.381 ± 0.081	4.0245 ± 0.0086	1.509 ± 0.034	1.69 ± 0.12	Andrievsky et al. (2002)

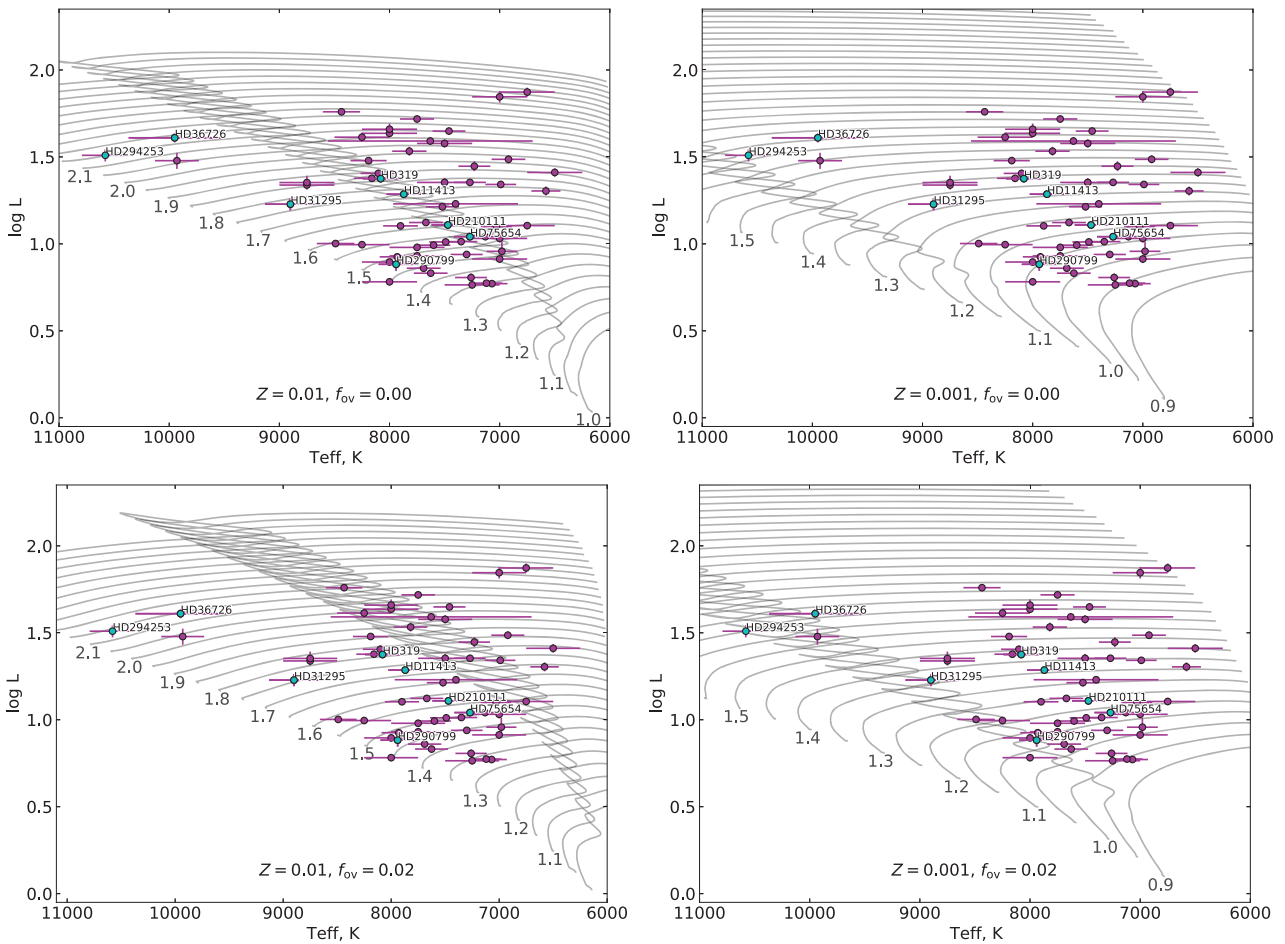


Figure 5. HR diagrams with stellar evolution tracks of different metallicity Z and overshooting f_{ov} . The positions of the 58 stars with reliable atmospheric parameters are shown with their 1σ uncertainties in magenta. The eight stars studied in detail in this Section 5 are also plotted in cyan and labelled by their HD numbers. Evolutionary models all use initial hydrogen abundance $X_{\text{ini}} = 0.7$, rotational velocity $V_{\text{rot}} = 60 \text{ km s}^{-1}$, and are shown at intervals of $0.05 M_{\odot}$. Mass labels in M_{\odot} are given at the ZAMS for every second track.

et al. 1994; Langanke & Martínez-Pinedo 2000). Screening was included via the prescriptions of Salpeter (1954), Dewitt, Graboske & Cooper (1973), Alastuey & Jancovici (1978), and Itoh et al. (1979). Thermal neutrino loss rates were from Itoh et al. (1996).

Fig. 5 shows that the stars can be found in a range of evolutionary stages, with the results depending strongly on metallicity and overshooting efficiency. For $Z = 0.010$ and $f_{\text{ov}} = 0.02$, most of

the 58 stars in the wider sample and all of the eight-star subsample are located on the main sequence. For the same metallicity but with no overshooting, only half of the wider sample and half of the eight-star subsample are definitely on the main sequence, while the others could be on the main sequence, in the post-main-sequence contraction phase, or after the loop and crossing the Hertzsprung Gap.

Table 5. Models of the eight target stars, fitting their observed values of effective temperature, luminosity, and rotational velocity. The model parameters given in the columns are: evolutionary stage of the star (MS = main sequence, CP = contraction phase, AL = after the loop), metallicity, initial hydrogen abundance, the extent of overshooting, initial rotational velocity on the zero-age main sequence, current rotational velocity, stellar mass, age, effective temperature, luminosity, radius, and surface gravity.

Star	ES	Z	X_{ini}	f_{ov}	$V_{\text{rot, ini}}$ (km s^{-1})	V_{rot} (km s^{-1})	M (M_{\odot})	star age (years)	R (R_{\odot})	log g
HD 319	MS	0.01	0.7	0.00	79	60	1.865	7.287×10^8	2.4578	3.9274
	CP	0.01	0.7	0.00	76	60	1.806	8.799×10^8	2.4577	3.9137
	AL	0.01	0.7	0.00	74	60	1.734	9.985×10^8	2.4578	3.8958
	MS	0.01	0.7	0.02	76	60	1.811	8.795×10^8	3.9096	3.9148
	AL	0.001	0.7	0.00	75	60	1.338	1.528×10^9	2.4577	3.7834
	AL	0.001	0.7	0.02	76	60	1.331	1.630×10^9	2.4577	3.7810
HD 11413	MS	0.010	0.7	0.00	147	125	1.782	8.081×10^8	2.3676	3.9403
	CP	0.010	0.7	0.00	138	125	1.707	1.034×10^9	2.3673	3.9217
	AL	0.010	0.7	0.00	136	125	1.662	1.122×10^9	2.3673	3.9099
	MS	0.010	0.7	0.02	142	125	1.736	9.605×10^8	2.3675	3.9288
	AL	0.001	0.7	0.00	128	125	1.264	1.858×10^9	2.3673	3.7912
	AL	0.001	0.7	0.02	131	125	1.267	1.896×10^9	2.3673	3.7923
HD 31295	MS	0.01	0.7	0.00	123	115	1.786	4.327×10^8	1.7318	4.2128
	MS	0.01	0.7	0.02	121	115	1.767	4.953×10^8	1.7318	4.2082
	AL	0.001	0.7	0.00	109	115	1.276	1.717×10^9	1.7318	4.0668
	MS	0.001	0.7	0.02	121	115	1.374	1.393×10^9	1.7319	4.0988
	CP	0.001	0.7	0.02	117	115	1.342	1.558×10^9	1.7318	4.0887
	AL	0.001	0.7	0.02	113	115	1.269	1.835×10^9	1.7319	4.0642
HD 36726	MS	0.01	0.7	0.00	94	80	2.189	3.450×10^8	2.1510	4.1128
	MS	0.01	0.7	0.02	92	80	2.149	4.075×10^8	2.1511	4.1047
	AL	0.001	0.7	0.00	95	80	1.613	8.367×10^8	2.1513	3.9802
	AL	0.001	0.7	0.02	94	80	1.573	1.017×10^9	2.1511	3.9693
HD 75654	MS	0.01	0.7	0.00	57	45	1.547	1.207×10^9	2.1275	3.9716
	CP	0.01	0.7	0.00	55	45	1.494	1.491×10^9	2.1276	3.9565
	AL	0.01	0.7	0.00	53	45	1.434	1.713×10^9	2.1277	3.9387
	MS	0.01	0.7	0.02	56	45	1.512	1.413×10^9	2.1276	3.9616
	AL	0.001	0.7	0.00	50	45	1.072	3.168×10^9	2.1276	3.8124
	AL	0.001	0.7	0.02	51	45	1.082	3.083×10^9	2.1276	3.8162
HD 210111	MS	0.01	0.7	0.00	40	30	1.613	1.081×10^9	2.2095	3.9570
	CP	0.01	0.7	0.00	39	30	1.565	1.301×10^9	2.2093	3.9439
	AL	0.01	0.7	0.00	38	30	1.494	1.527×10^9	2.2094	3.9237
	MS	0.01	0.7	0.02	39	30	1.573	1.281×10^9	2.2093	3.9461
	AL	0.001	0.7	0.00	37	30	1.123	2.710×10^9	2.2093	3.7999
	AL	0.001	0.7	0.02	38	30	1.134	2.637×10^9	2.2092	3.8041
HD 290799	MS	0.01	0.7	0.00	71	70	1.472	3.401×10^8	1.3910	4.3191
	MS	0.01	0.7	0.02	71	70	1.468	3.791×10^8	1.3909	4.3180
	AL	0.001	0.7	0.00	64	70	1.051	3.063×10^9	1.3910	4.1727
	MS	0.001	0.7	0.02	70	70	1.116	2.423×10^9	1.3910	4.1990
	CP	0.001	0.7	0.02	66	70	1.073	2.939×10^9	1.3910	4.1819
	AL	0.001	0.7	0.02	64	70	1.038	3.257×10^9	1.3911	4.1674
HD 294253	MS	0.01	0.7	0.00	72	70	2.120	9.592×10^6	1.5180	4.4018
	MS	0.01	0.7	0.02	72	70	2.120	1.188×10^7	1.5180	4.4016
	MS	0.001	0.7	0.00	80	70	1.682	5.818×10^8	1.4945	4.3148
	CP	0.001	0.7	0.00	75	70	1.611	7.521×10^8	1.4944	4.2960
	AL	0.001	0.7	0.00	74	70	1.570	8.097×10^8	1.4944	4.2848
	MS	0.001	0.7	0.02	77	70	1.640	6.912×10^8	1.4944	4.3038

For low metallicity ($Z = 0.001$) and without overshooting from the convective core, all stars but one are beyond the main sequence. The only exception, HD 294253, is ambiguously located in one of the three different stages. If we add overshooting with $f_{\text{ov}} = 0.02$, the majority of the wider sample are crossing the Hertzsprung Gap and the rest are ambiguously located.

From the evolutionary modelling it follows that if the λ Boo stars have intrinsically low metallicity of the order of -1 dex, then most of them must have evolved beyond the main sequence. Arguments concerning the time-scales of the evolutionary phases show this to be unlikely. Using the evolutionary models, we calculated the time taken for a $1.7 M_{\odot}$ star to evolve 100 K (between 7100 and

7000 K) along its evolutionary track at the main-sequence, the contraction phase and the after-the-loop phase. These durations are 25.8, 1.6, and 0.7 Myr, respectively. Since the stars are almost equally luminous in these phases, no Malmquist bias is present. It can be surmised that the stars are 16 to 37 times more likely to be on the main-sequence than in the other phases. The evolutionary models present a strong argument that the stars are not globally metal poor.

It is also known from TESS data that several λ Boo stars have pulsation properties consistent with young main-sequence ages (Bedding et al. 2020; Appendix A, here). We conclude with high confidence that λ Boo stars are only superficially metal-weak.

5.2 Pulsation models

Another way to test whether λ Boo stars are globally metal weak is to construct pulsation models using two metallicities ($Z = 0.01$ and $Z = 0.001$) and to see which is a better fit to the modes observed with TESS. We used the eight-star subsample (Table 4) for this, constructing models that fit their observed T_{eff} , L , and $v \sin i$.

The model parameters are given in Table 5. In all but one case we considered models within the corresponding error box. The only exception was HD 294253 when calculated with $Z = 0.010$, because at this metallicity star lies beneath the ZAMS. In that one case we adopted a slightly lower value of the effective temperature.

Depending on the values of metallicity and overshooting, a given star can be located in one of the three aforementioned evolutionary phases. Hence, by varying metallicity and overshooting, a given HR position based on T_{eff} and L results in quite different model parameters such as mass and age (Table 5). Therefore, it is impossible to independently constrain the masses and ages of the stars. The value of stellar radius, however, is insensitive to the metallicity, overshooting and mass because it is determined from the T_{eff} and L alone. The radii thereby determined for the eight target stars are given in Table 4.

In general, the models suggest slightly higher surface gravities than the spectroscopic values, although most agree within the typical uncertainty of 0.2 dex. There are two exceptions, HD 290799 and HD 294253. Their models have lower $\log g$ than the observations, but the observed values of $\log g = 4.5$ are unphysical because they place the star beneath the main sequence at $Z = 0.01$.

The pulsation properties are shown in Figs 6–13, where we plot the instability parameter, η , as a function of frequency for all considered stellar models. The instability parameter is a measure of the energy gained/lost during the pulsation cycle, normalized to the range $[-1, 1]$ (Stellingwerf 1978). If $\eta > 0$, then the corresponding mode is excited in a model. In our calculations we used the customized non-adiabatic pulsation code of Dziembowski (1977). While models with and without overshooting were calculated, we found that overshooting affected only the possible evolutionary stages (as described in Section 5.1), with negligible difference to the excitation (Fig. 6). Hence from Fig. 7 onwards, only the models with overshooting are shown. Although helium is predicted to be overabundant in λ Boo stars, there are no helium lines in A stars to use for helium abundance measurements. We did not enhance the helium abundance in our models (Figs 6–13), which would have increased the pulsational driving, but on a case-by-case basis we do examine the driving effect of helium enhancement.

5.2.1 Pulsation model analysis

5.2.2 HD 319

The pulsation models of HD 319 predict no excited modes (Fig. 6). The highest value of the instability parameter, η , is about -0.1 for modes with frequencies $15\text{--}20\text{ d}^{-1}$. The evolutionary phase only has an important impact on the low frequencies (g modes) but in no model are these excited. These results are consistent with the observations (Fig. 6, Table 1, and Appendices A and B), where no variability was seen.

5.2.3 HD 11413

Models calculated with $Z = 0.01$ sit just below the threshold for excitation, with $\eta \sim -0.001$ for frequencies of $\sim 20\text{ d}^{-1}$ (Fig. 7). For

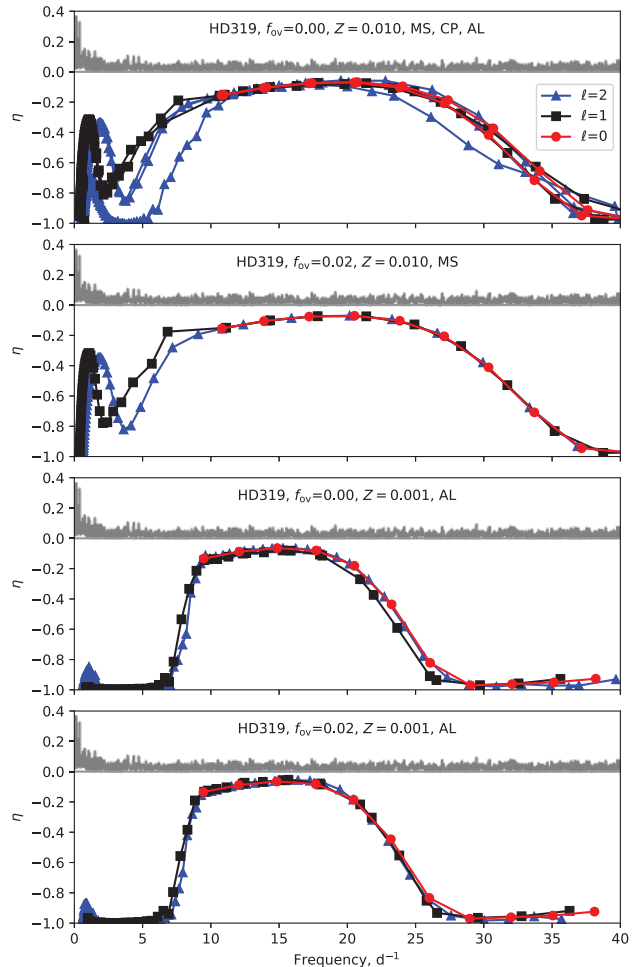


Figure 6. Instability parameter, η , as a function of the pulsation frequency for the seismic models of HD 319. Mode driving exceeds damping and modes are excited where $\eta > 0$, but no modes are observed or predicted to be excited in HD 319. Mode degrees $\ell = 0, 1$, and 2 were considered. The top panel shows models calculated with $Z = 0.010$ and $f_{\text{ov}} = 0.00$. In the next panel we show the overshooting effect with $f_{\text{ov}} = 0.02$. In the bottom two panels the models are calculated with lower metallicity ($Z = 0.001$) and two values of f_{ov} (0.00 and 0.02). All possible evolutionary stages were considered (MS = main sequence, CP = contraction phase, AL = after the loop).

lower metallicity, $Z = 0.001$, the models are much more evolved, after the loop, and radial and dipole modes with frequencies of $10\text{--}15\text{ d}^{-1}$ reach $\eta \sim 0.012$. The high frequency modes are very effectively damped in these post-main-sequence models, so the observed frequencies, which mostly lie in the range $20\text{--}30\text{ d}^{-1}$, are slightly better matched by the higher metallicity models, even though η is slightly negative in the latter. We note that an increase in the helium abundance by about 0.1 produces excited modes in the observed frequency range in the higher metallicity models.

5.2.4 HD 31295

The models do not predict any excited modes, regardless of metallicity, overshooting, and evolutionary phase (Fig. 8). This is in contradiction with the observations, which indicate pulsations with high frequencies of the order of 40 d^{-1} . Interestingly, in this frequency region the instability coefficient η does have a maximum. However, an increase of the helium abundance to $Y = 0.39$ is needed

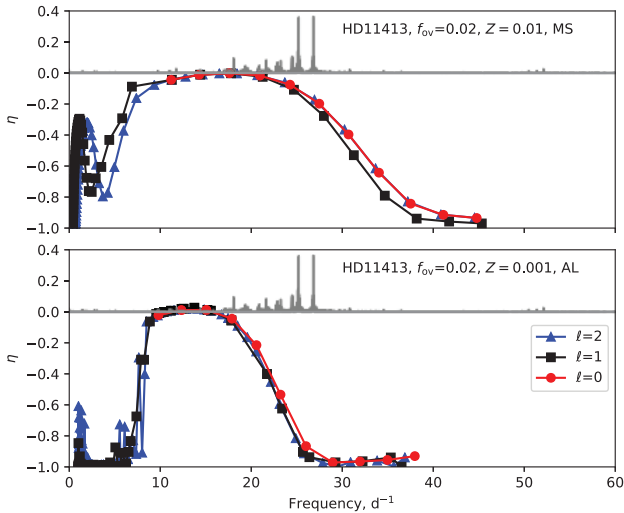


Figure 7. The same as in Fig. 6, but for HD 11413 and showing only the models with overshooting.

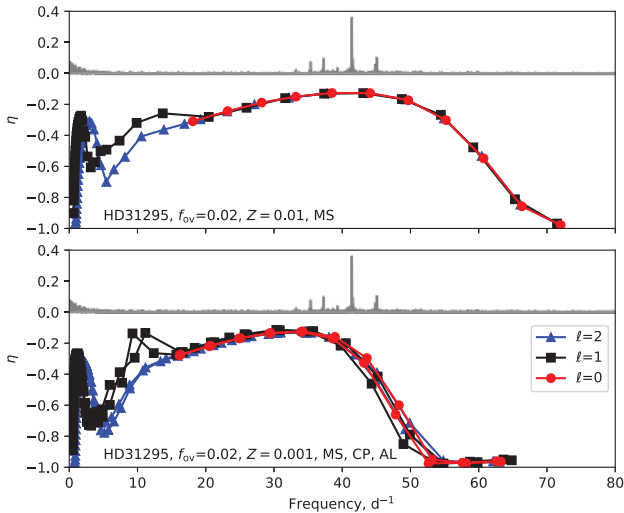


Figure 8. The same as in Fig. 7, but for HD 31295.

in order to excite modes at $\sim 40 \text{ d}^{-1}$. The higher metallicity model is slightly preferred because it has η reaching a maximum closer to the observed frequency of maximum power.

5.2.5 HD 36726

Neither the models nor the observations show excited modes (Fig. 9). The instability parameter is the highest for g modes, but its value is still no greater than -0.2 .

5.2.6 HD 75654

A very clear instability occurs for all considered models of HD 75654 (Fig. 10). Models calculated with $Z = 0.010$ have unstable modes with frequencies from about 7 to 45 d^{-1} . For lower metallicity, $Z = 0.001$, the unstable modes occur over a smaller frequency range ($10\text{--}25 \text{ d}^{-1}$). Our models are computed with frozen convection, which is not strictly valid for high oscillation frequencies or for stars like this where as much as 90 per cent of the flux is transported by convection. Nonetheless, the observed frequencies of the order of 14 d^{-1} are consistent with all inspected models. There

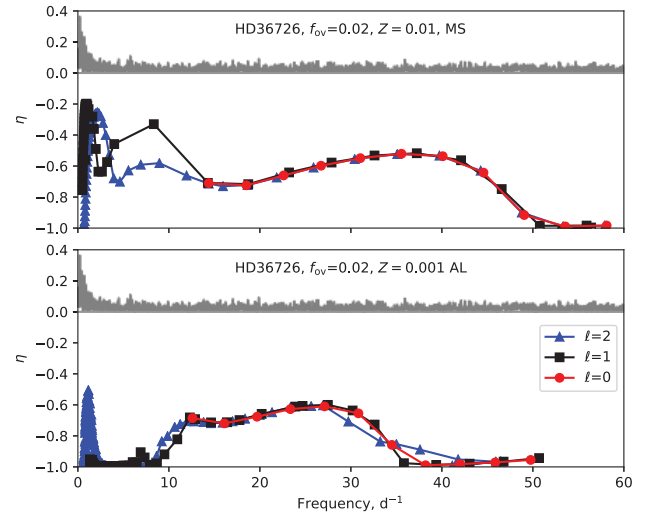


Figure 9. The same as in Fig. 7, but for HD 36726.

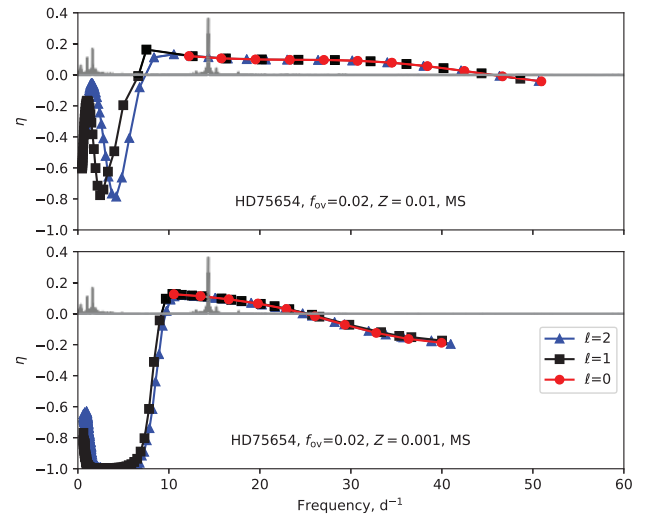


Figure 10. The same as in Fig. 7, but for HD 75654.

is no clear preference for metallicity, since the higher metallicity overpredicts instability at higher frequencies ($\sim 30 \text{ d}^{-1}$) but also comes closer to describing the observed lower frequencies ($\sim 1 \text{ d}^{-1}$).

5.2.7 HD 210111

Models predict unstable modes between ~ 7 and 30 d^{-1} for $Z = 0.010$, and from ~ 10 to 20 d^{-1} for $Z = 0.001$ (Fig. 11). These results agree well with the high frequencies seen in the TESS data (see also Appendix A). The detected low frequencies cannot be explained by post-main-sequence models since they are strongly damped ($\eta = -0.8$) in the advanced evolution phases. They are also damped on the main sequence, but with $\eta \sim -0.2$. The higher metallicity model is a better match to both the higher and lower frequencies.

5.2.8 HD 290799

The instability parameter is slightly positive for this star (Fig. 12). The excited modes are at frequencies in the range $15\text{--}40 \text{ d}^{-1}$ for higher metallicity models and $15\text{--}30 \text{ d}^{-1}$ for lower metallicity models. TESS observations show pulsation frequencies in a wide

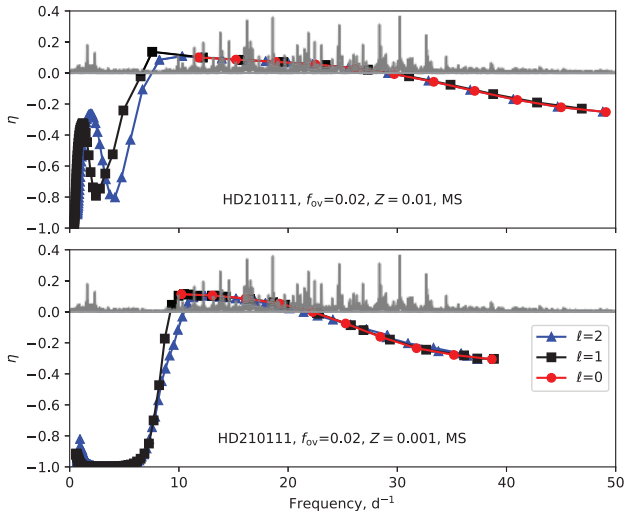


Figure 11. The same as in Fig. 7, but for HD 210111.

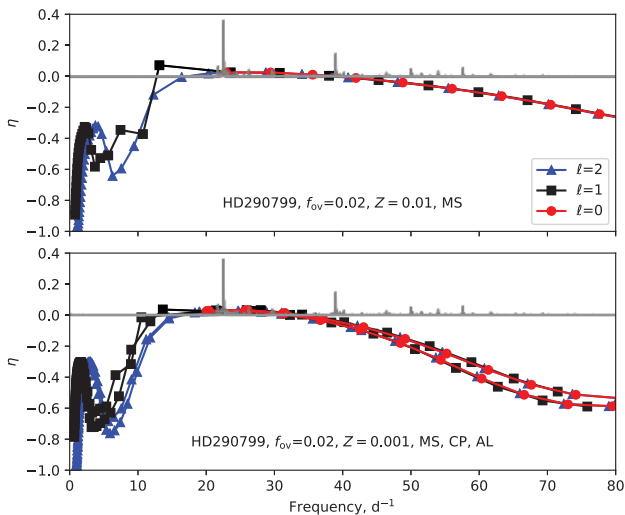


Figure 12. The same as in Fig. 7, but for HD 290799.

range of $20\text{--}60\text{d}^{-1}$. The high frequencies fit better to the main sequence models and a preference towards high metallicity was found. However, caution is again required by the use of the frozen convection approximation. It is noteworthy that Bedding et al. (2020) found a very high $\Delta\nu$ of 7.55d^{-1} , suggesting this star is very young (see Appendix A).

5.2.9 HD 294253

This is the most massive star we analysed. Its models predict no unstable modes, regardless of the chemical composition and evolution phase (Fig. 13). No variability was observed with TESS, either (Table 1, Appendices A and B).

5.2.10 Pulsation model summary

Three of the eight stars for which we constructed pulsation models are non-pulsators. The pulsation models accurately reproduce this. In three of the remaining five targets, excited modes were reproduced by the models. Where any metallicity was preferred, it was always the higher metallicity ($Z = 0.01$) model, though HD 210111 is the only target for which this case can be made strongly. While the

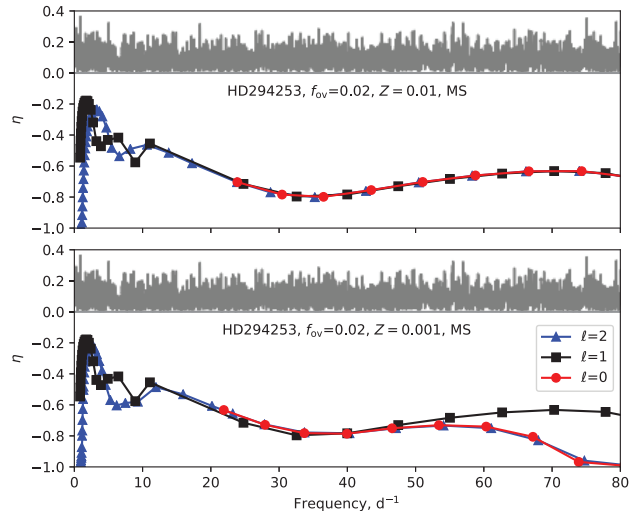


Figure 13. The same as in Fig. 7, but for HD 294253.

frozen convection approximation and the simplicity of the analysis are no substitute for dedicated modelling, we conclude that such analyses are valuable in evaluating the global metallicity of λ Boo stars, which in this case has hinted that these λ Boo stars are unlikely to be metal-poor throughout.

6 CONCLUSIONS

We analysed the TESS light curves of 70 southern λ Boo stars to evaluate whether they pulsate as δ Sct stars. We discovered four objects to be binaries: HD 31508 and HD 94390 are heartbeat stars, and HD 168947 and HD 74423 are eclipsing binaries. Sixteen of the studied λ Boo stars are not δ Sct stars at the precision of TESS data, although many of these lie far from the δ Sct instability strip, where pulsation is not expected. By reference to a sample of 1826 *Kepler* δ Sct stars, we computed the expected pulsator fraction, given the location of stars with respect to the instability strip. We found that the fraction of λ Boo stars that show δ Sct pulsation is about twice that of normal stars. This has important implications: firstly, having more pulsating λ Boo stars opens up more possibility for detailed asteroseismic modelling; and secondly, it suggests that the mechanism that produces the chemical peculiarity may also make the stars more susceptible to pulsation. Indeed, this link was already hypothesized (Murphy 2014).

As a step towards asteroseismology of the λ Boo stars in this sample, we have attempted to identify the fundamental mode based on the period–luminosity relationship, échelle diagrams, and Petersen diagrams (period ratios). We identified the fundamental mode in 20 of the 40 pulsating λ Boo stars that have 2-min TESS data and reliable atmospheric parameters. In an additional eight stars, we can conclude that the stars are not pulsating in the fundamental mode, while in the remaining 12 it was not possible to identify the fundamental mode unambiguously. We confirmed the finding by Bedding et al. (2020) that some λ Boo stars have high-frequency pulsations with remarkably regular frequency patterns. These patterns facilitate mode identification, allowing radial ($\ell = 0$) and dipole ($\ell = 1$) modes to be distinguished, and the radial order (n) to be inferred. This enables a precise stellar age to be calculated and supports the hypothesis that λ Boo stars are young (e.g. Kama et al. 2015).

An important question is whether λ Boo stars are globally metal-poor, or just metal-poor at the surface. It was hoped that aster-

oseismology would answer this question, but the first major step towards this was at odds with other analysis methods. Specifically, the asteroseismic investigation of 29 Cyg – the first λ Boo star discovered to be variable (Gies & Percy 1977) – suggested a global metal weakness (Casas et al. 2009), whereas evolution models for several other λ Boo stars suggested they were of global solar abundance (Pauzen et al. 2015). We used *Gaia* DR2 parallaxes to provide precise localizations on the HR diagram for the λ Boo stars. By comparing with stellar evolution models with a range of metallicities, we found it implausible that λ Boo stars are globally metal poor. The most likely conclusion is that they are all main-sequence stars of approximately solar metallicity. Otherwise, if the λ Boo stars of our sample are metal-poor throughout, the majority of them would have to be in the extremely short-lived phase of evolution after the main-sequence, where stars cross the Hertzsprung gap.

We also calculated pulsation models for a subset of eight stars. We found that the observed pulsations tend to be better matched to the higher metallicity models ($Z = 0.01$) than the low metallicity models ($Z = 0.001$), albeit with small-number statistics. However, the simplicity of the models leaves much room for refinement. Future modelling could use time-dependent convection, incorporate turbulent convection into the driving (Houdek 2000; Antoci et al. 2014), use a higher helium abundance, and be based on mode identification as far as possible. We provided a start for this mode identification, and physical parameters on which to base future pulsation models. This opportunity is an exciting prospect for λ Boo star research.

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SUPPORTING INFORMATION

Supplementary data are available at *MNRAS* online.

- Figure S1** Fourier transforms of λ Boo star light curves.
- Figure S2** Fourier transforms of λ Boo star light curves.
- Figure S3** Fourier transforms of λ Boo star light curves.
- Figure S4** Fourier transforms of λ Boo star light curves.
- Figure S5** Fourier transforms of λ Boo star light curves.
- Figure S6** Fourier transforms of λ Boo star light curves.
- Figure S7** Fourier transforms of λ Boo star light curves.
- Figure S8** Fourier transforms of λ Boo star light curves.
- Figure S9** Fourier transforms of λ Boo star light curves.
- Figure S10** Fourier transforms of λ Boo star light curves.
- Figure S11** Fourier transforms of λ Boo star light curves.
- Figure S12** Fourier transforms of λ Boo star light curves.
- Figure S13** Fourier transforms of λ Boo star light curves.
- Figure S14** Fourier transforms of λ Boo star light curves.

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APPENDIX A: NOTES ON THE FOURIER TRANSFORMS OF INDIVIDUAL STELLAR LIGHT CURVES

TESS has made possible a much broader investigation of pulsating λ Boo stars than the six already known (Table A1). In this appendix, we describe the features of TESS light curves (and their Fourier transforms) of southern λ Boo stars.

HD 319, sector 2

This non- δ Sct star is also discussed in Section 5. No p modes are evident at the 0.01 mmag level, even though the star is located within the instability strip. The Fourier transform features peaks at low frequency ($<0.5 \text{ d}^{-1}$), with amplitudes <0.30 mmag, which could be rotational or instrumental in origin.

HD 3922, sector 1–2

The low frequencies ($<1 \text{ d}^{-1}$) are caused by trends in the data near the sector boundaries, and are not pulsational in origin. The p modes are unaffected by this. The star does not appear to be a γ Dor variable, unless pulsation is masked by the low-frequency trends at the $\sim 0.3 \text{ mmag}$ level. This star lies near the TAMS, which explains the low frequencies of the strongest p modes ($5\text{--}10 \text{ d}^{-1}$). Identifying these modes is difficult. f_1 is a good candidate for the first overtone mode because the period ratio with the peak at 11.15 d^{-1} coincides with the expected one for the second radial overtone mode. No peak clearly lines up with the predicted frequency of the fundamental mode from the P–L relation (7.91 d^{-1}), or from frequency ratios using the three strongest modes.

HD 4158, sector 3

Period ratios suggest that $f_2 = 7.106 \text{ d}^{-1}$ is the fundamental mode and $f_3 = 9.169$ is the first radial overtone (the ratio $f_2/f_3 = 0.775$). In addition to the p modes, there are also peaks at 0.57 and 1.70 d^{-1} , with amplitudes of 0.2 and 0.3 mmag . These could be r and g modes, respectively. The similarity of the p-mode spectrum to that of HD 3922 is noteworthy – the two stars seem to be equally evolved but HD 4158 is $\sim 500 \text{ K}$ hotter.

HD 6870, sectors 1–3

The pulsations are dominated by three p modes with amplitudes between 1 and 2 mmag , but there are also groups of peaks between 1.1 and 4.6 d^{-1} that are likely g modes and perhaps prograde sectoral modes. No pair from the three dominant peaks has a frequency ratio near 0.77 that would correspond to the fundamental mode and first radial overtone, but the strongest (and highest frequency) of the three, $f_1 = 17.87 \text{ d}^{-1}$, lies very close to the predicted value from the P–L relation of 17.80 d^{-1} .

HD 7908, sector 3

Not a δ Sct star at the 0.02 mmag level. Broad power excesses at $1.0\text{--}1.5 \text{ d}^{-1}$ and at $2.4\text{--}2.9 \text{ d}^{-1}$ suggest this is a γ Dor star. It lies just inside the red edge of the theoretical δ Sct instability strip (Dupret et al. 2005) but outside the observed one (Murphy et al. 2019).

Table A1. Known pulsation properties of λ Boo stars from the literature. Only objects with long time-series are included.

HD	Frequency d^{-1}	Ampl. mmag	Filter	Reference
11413	25.206	9.1	V	Koen et al. (2003)
	24.099	5.3	–	–
	24.493	5.0	–	–
83041	14.5293	1.6	SWASP	Paunzen et al. (2015)
111786	31.01	8.0	V	Paunzen et al. (1998)
	20.31	6.0	–	–
	33.36	5.8	–	–
184779	12.5687	13.9	SWASP	Paunzen et al. (2015)
	13.8351	6.8	–	–
	10.4756	3.0	–	–
210111	28.47	9.5	V	Breger et al. (2006)
	20.81	2.6	–	–
	18.59	2.6	–	–
290799	22.5300	7.2	SWASP	Paunzen et al. (2015)

HD 11413, sectors 2–3

Pulsation models of this star are presented in Section 5. We measure a characteristic spacing for the p modes of $3.6\text{--}3.8 \text{ d}^{-1}$ that could be the large spacing, $\Delta\nu$ (see Bedding et al. 2020). For a spacing of 3.8 d^{-1} , the fundamental mode is probably the peak at 11.8 d^{-1} , but there is not independent support from the P–L relation for this. Period ratios suggest the peaks at 14.424 and 18.06 d^{-1} could be the first and second radial overtone modes (measured ratio = 0.799), but these are not compatible with the fundamental mode measured from the échelle. The peaks near 50 d^{-1} are harmonics and combinations of the strong peaks.

HD 13755, sector 3

The Fourier spectrum is quite dense with p modes at relatively low frequencies ($5\text{--}15 \text{ d}^{-1}$). The predicted fundamental mode frequency is 8.00 d^{-1} , which is close to f_1 at 7.93 d^{-1} . It then follows from period ratios that $f_2 = 12.66 \text{ d}^{-1}$ is probably the second radial overtone, which we predict to lie at $\sim 12.75 \text{ d}^{-1}$. No g modes are obvious.

HD 17341, sectors 3–4

The p modes span a broad frequency range. This star lies on the ZAMS and has a measured $\Delta\nu = 5.90$ (Bedding et al. 2020). The P–L relation suggests a fundamental mode frequency of 16.3 d^{-1} , which does fall close to an observed but weak peak at 16.4 d^{-1} . However, there are other weak peaks at similar frequencies and the échelle diagram shows no clear candidate for the fundamental mode (Fig. A1). Some low-frequency peaks are evident and could be g-mode oscillations.

HD 23392, sector 4

Not a δ Sct star at the 0.02 mmag level, which is unsurprising given its location well beyond the blue edge of the δ Sct instability strip. The star does not appear to pulsate in g modes, either. The light-curve drifts on time-scales of one sector, dipping towards sector boundaries, so the peaks at low frequency are largely instrumental.

HD 28490, sector 5

The oscillation frequencies are at the low end of the range expected for δ Sct p modes, consistent with the star’s location near the TAMS. None of the dominant frequencies shows the expected period ratios for radial modes. Despite the high luminosity, an échelle diagram does show a ridge (using $\Delta\nu = 2.32 \text{ d}^{-1}$), which implies the peak at $f_2 = 6.95 \text{ d}^{-1}$ is the fundamental mode. This identification is consistent with the P–L relation. The stronger peak at lower frequency ($f_1 = 5.2 \text{ d}^{-1}$) is presumably a mixed mode. There are no clear low frequencies (i.e. no g modes).

HD 28548, sector 5

The Fourier transform shows many almost equally spaced pulsation frequencies at high frequency ($\sim 35\text{--}65 \text{ d}^{-1}$). Bedding et al. (2020) found a very high $\Delta\nu = 7.56 \text{ d}^{-1}$, consistent with the star’s location near the ZAMS. They also matched it to a model, which had an age of 270 Myr . The fundamental mode is weak or missing. Two clear peaks exist at low frequency, but are not harmonics of each

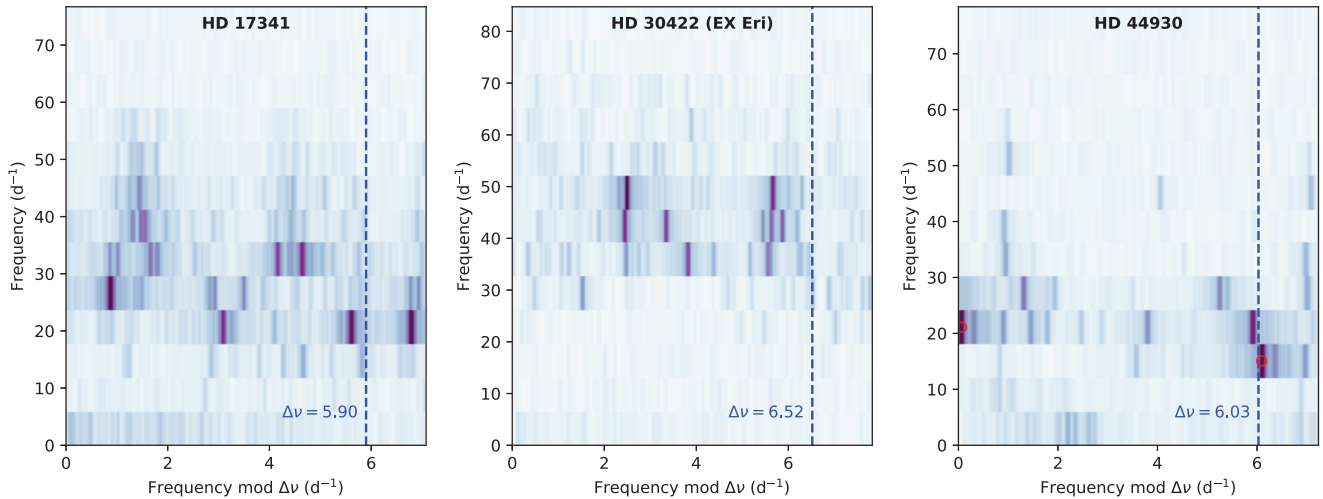


Figure A1. Échelle diagrams of HD 17341, HD 30422, and HD 44930. While HD 17341 shows two nearly vertical ridges at $x \approx 1.5$ and $x \approx 4.3 \text{ d}^{-1}$, there is no clear candidate for the fundamental mode. By contrast, the fundamental mode is clear in HD 44930 (red circle).

other. Instead, they are themselves the first two peaks of a harmonic series. The harmonic nature argues against a γ Dor classification, but the cause is currently undetermined.

HD 30422, sector 5

Also known as EX Eri. The high-frequency p modes exhibit a regular spacing, with a large spacing of 6.52 d^{-1} . This is one of the exemplars from Bedding et al. (2020). The échelle suggests this star does not pulsate in the fundamental mode (Fig A1).

HD 31295, sector 5

Pulsation models of this star are presented in Section 5. Its p modes are at relatively high frequency ($\sim 40 \text{ d}^{-1}$). There are few of them, resulting in a simple pulsation spectrum, and no excited modes lie near the predicted fundamental mode. The oscillation spectrum of this hot star seems to be incompatible with the period inferred from the P–L relation, which is probably because of the absence of a colour term. There are no clear g modes.

HD 31508, sectors 1–13

This target falls on silicon for sectors 1–13, only a subset of which has 2-min data available (2, 3, 6, 12, 13). We analysed the 2-min data only. The target is a heartbeat star (an eccentric ellipsoidal variable) with a period of 19.2 d. The periastron brightening is stronger than the dip immediately preceding it (Fig. A2).

HD 36726, sector 6

This non-pulsator is also discussed in Section 5. No pulsation is evident down to a grass level of 0.025–0.030 mmag. The star is much hotter than the blue edge of the δ Sct instability strip.

HD 37411, sector 6

This FFI-only target is not a δ Sct star at the 0.2 mmag level. It does, however, show non-periodic excursions of ~ 25 mmag (peak-to-peak) on time-scales of days. This variability is likely

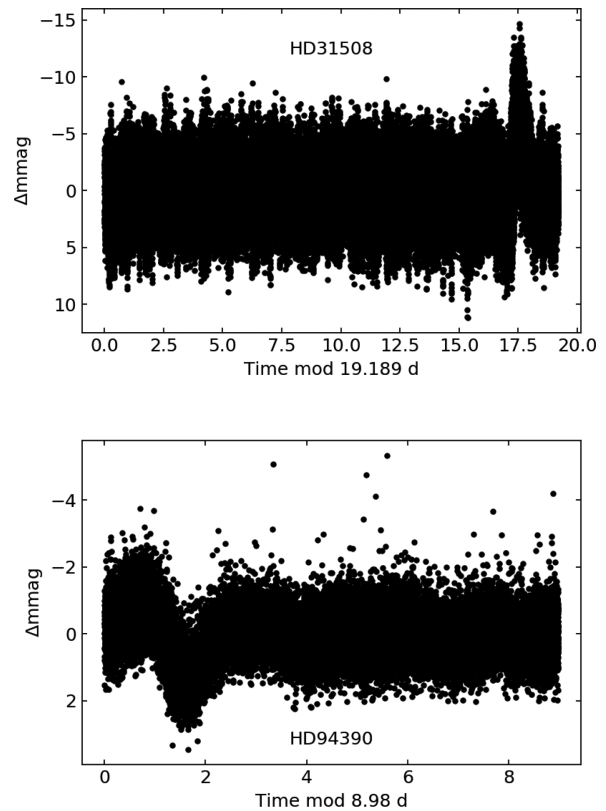


Figure A2. TESS light curves of the λ Boo heartbeat stars HD 31508 and HD 94390, folded on their orbital periods of 19.189 and 8.98 d, respectively.

associated with its pre-MS nature – McNamara & Huels (1983) found HD 37411 to be a member of the Orion Nebula Trapezium cluster with 96 per cent probability and Gray & Corbally (1998) studied the spectrum of this star in detail, finding emission lines and a variable spectrum. It has near- and far-infrared excesses and is accreting matter (Grady et al. 1996; Gray & Corbally 1998; Malfait, Bogaert & Waelkens 1998). According to the star’s spectral type (hA2 Vae kA8mA8 λ Boo; Murphy et al. 2015) our T_{eff} of $\sim 11400 \text{ K}$

is probably too hot. Either way, HD 37411 lies beyond the blue edge of the δ Sct instability strip, so its lack of pulsation is unsurprising. We also find a large extinction of $A_V = 1.56$ mag for this target.

HD 38043, sectors 1–3, 5–6, 8–13

The TESS light curve spans 357 d and includes all sectors except 4 and 7. The data are at 2-min cadence in all available sectors. The p modes appear to be dominated by two nearly equally spaced multiplets: the first contains the strongest three modes – the same ones for which frequencies and amplitudes are provided in Table 1 – with a spacing of ~ 0.51 d $^{-1}$. The second is slightly lower in frequency and amplitude and has a spacing of ~ 0.55 d $^{-1}$. These could be rotationally split multiplets of modes with similar but non-identical n and ℓ . This star is highly multiperiodic at lower frequencies, consistent with g-mode pulsation, though we could not identify any clear period spacings.

HD 41958, sectors 1, 4–5, 7–8, 11

The three strongest p modes have similar amplitudes (7.5–8.2 mmag). They do not have any obvious harmonic relationship, nor do they have the frequency ratio (0.77–0.78) expected of the fundamental mode and first radial overtone modes.

HD 42503, sectors 5–7

The mutli-periodic nature of the light curve at frequencies near 5 d $^{-1}$ is rather unusual. The frequency content does not strongly resemble that of γ Dor stars, but is also quite different from typical δ Sct stars. No radial modes could be identified from frequency ratios of the stronger peaks. The peaks at 0.99 and 2.00 d $^{-1}$ are not perfect harmonics, and another peak at 3.94 d $^{-1}$ also falls on this near-harmonic pattern.

HD 43533, sectors 6–7

The star is a γ Dor variable and the light curve is dominated by low-frequency variability, with dozens of significant peaks below 6 d $^{-1}$, and this variability spills into the δ Sct frequency domain. Excluding this variability, which presumably does not include independent p modes, there are only a few peaks in the Fourier transform that make this a δ Sct star. All three extracted peaks are significant, with SNR > 4, despite their low amplitudes. The low amplitudes perhaps suggest these peaks belong to a contaminating source, but the *Gaia* DR2 data show no likely contaminant. It is noteworthy that this star is located below the ZAMS compared to tracks of solar metallicity.

HD 44930, sector 5

The Fourier transform is dominated by several p modes of mmag amplitude. A group of weaker peaks with frequencies between 2 and 3 d $^{-1}$ might be g modes. The absence of peaks at even lower frequencies would then suggest these are $\ell = 1$, $m = 1$ prograde modes, which suggests the near-core region rotates rapidly. This star is in the Octans association and has a measured $\Delta\nu$ of 6.03 d $^{-1}$ (Bedding et al. 2020). The échelle diagram (Fig A1) suggests $f_1 = 18.15$ d $^{-1}$ is the fundamental mode, even though this is not in agreement with the P–L prediction.

HD 47425, sectors 6–7

The two dominant p modes are too close in frequency to be the fundamental mode and first radial overtone. Indeed, it is not possible to clearly identify radial modes by period ratios in this star. There are many p modes with mmag amplitudes, though no g modes are evident at amplitudes similar to the p modes.

HD 46722, sectors 6–7

This δ Sct star has a semiregular series of frequencies from ~ 30 –50 d $^{-1}$ with large spacing $\Delta\nu = 6.45$ d $^{-1}$ (Bedding et al. 2020). Three humps of peaks at lower frequency suggest this is also a γ Dor star. The échelle diagram and the P–L relation suggest that $f_1 = 20.76$ d $^{-1}$ is the fundamental mode.

HD 68695, sectors 7–8

In these FFI data, this δ Sct star has only three significant p-mode peaks. The identified oscillation frequencies may be Nyquist aliases – one of the aliases of the identified f_1 has a higher amplitude by 1.9 μ mag, which may suggest the ‘alias’ is in fact the real peak (see Murphy 2015). However, the noise level is an order of magnitude higher than this and the star is not located near the ZAMS (where oscillation frequencies tend to be higher), so with the current observations, we suggest the given frequencies are more likely to be the real ones.

HD 73211, sector 8

The peak between the two strongest peaks is $f_4 = 19.25$ d $^{-1}$; the third strongest peak falls outside of the range of the Fourier plot in Appendix B. The triplet formed from f_1 , f_2 , and f_4 is not equally spaced, and none of these peaks falls near the predicted fundamental mode from the P–L relation.

HD 74423, sectors 9–11

The analysis of the TESS light curve of this fascinating λ Boo star is described in a dedicated paper (Handler et al. 2020). The entire variability can be represented with two independent frequencies and as many as 21 combination frequencies of amplitude > 0.01 mmag. The two independent frequencies are a binary orbital frequency, $f_{\text{orb}} = 0.63262$ d $^{-1}$, and a single p-mode oscillation at $f_1 = 8.7569$ d $^{-1}$. Handler et al. (2020) interpret the light curve as pulsations trapped in a single hemisphere of the star by tidal forces. The reader is referred there for more information.

HD 75654, sector 8–9

Pulsation models of this star are presented in Section 5. The strongest peak lies almost on top of the predicted fundamental mode frequency from the P–L relation. In addition to the p modes, there are also peaks at 0.36, 1.07, and 1.62 d $^{-1}$. These are not obvious harmonics and combinations, and may be r and g modes.

HD 76097, sector 8

The third peak falls close to the strongest one, and their spectral windows are not well separated. The frequencies for this star were therefore extracted by simultaneous non-linear least-squares fitting, rather than by the automated algorithm. The fundamental mode

cannot be unambiguously identified because of the proximity of these strong peaks, and because there are no clear candidates for the first and second radial overtone modes for reference.

HD 80426, sector 8

The star is dominated by a 10-mmag peak at $f_1 = 11.01 \text{ d}^{-1}$. Despite first appearances, the next strongest peaks are not harmonics of this frequency. The ragged peak shape is simply the result of the spectral window. Since this star is located above the TAMS and at the red edge of the instability strip, this frequency appears to be too high to be the fundamental mode.

HD 83041, sectors 8–9

In addition to the p modes, there are many peaks below 5 d^{-1} , which suggests this star is also a γ Dor variable. The strongest peak agrees with the one identified in SuperWASP data by Paunzen et al. (2015).

HD 84159, sector 8

The broad peaks in the Fourier transform are simply due to a broad spectral window, caused by the large data gap between orbits. The p-mode amplitudes are very high. A pair of peaks near 1.8 d^{-1} could be prograde g modes. The group of peaks near 18 d^{-1} appear to be combination frequencies involving the strongest peak, f_1 , which is to be expected given the strong pulsation amplitudes. No frequency pair gives a period ratio consistent with both being radial modes.

HD 88554, sectors 9–10

An unremarkable low amplitude δ Sct star.

HD 94326, sectors 9–10

The light-curve morphology gradually changes from the start of sector 9 to the end of sector 10, with the amplitude of the low-frequency variability decreasing. Independent of this, the star appears to be a γ Dor variable, with mmag peaks between 0.8 and 1.4 d^{-1} . The strongest δ Sct peak is at 10.1 d^{-1} , and there is significant variability at many frequencies between this and the g modes at the ~ 0.3 mmag level. No peak stands out as the fundamental mode.

HD 94390, sectors 9–10

This is a heartbeat star (an eccentric ellipsoidal variable) with an orbital period of 8.98 d, perhaps exhibiting tidally excited oscillations. However, it is not a δ Sct star at the 0.02 mmag level.

HD 98069, sector 9

This object has both K2 and TESS (FFI) data. There has been remarkable change in pulsation amplitudes between the two data sets, beyond that which is explained by the change in passband. While the strongest mode is still at the same frequency, the relative amplitudes of other modes compared to this peak are very different. The second and third strongest modes in the K2 data are only marginal in the TESS data set. The extracted peaks given here are the TESS ones, for consistency with the other stars. We assume that the real peaks fall between 0 and the Nyquist frequency of 24 d^{-1} .

The temperature and luminosity of the star place it near the TAMS, so this assumption appears to be justified.

HD 100546, sector 11

The TESS light curve is unremarkable in most regards, and show this is not a δ Sct star. It is also too hot to lie in the δ Sct instability strip. Some low-frequency variability is evident, even against the rising background at low frequency. There could be g modes, or this could be attributed to rotation, or since this star is known to be actively accreting material (Vieira et al. 1999), the low frequencies could be accretion-induced variability.

HD 101412, sectors 10–11

Only FFI data are available in each sector. Only one peak is (questionably) evident in the Fourier transform, at 10.679 d^{-1} , but it has an SNR of 3.3, and we therefore do not consider it significant. Hence, HD 101412 is not a δ Sct star, despite being inside the instability strip.

HD 102541, sector 10

A known δ Sct star (V1023 Cen). The Fourier transform shows a series of oscillation frequencies that continue to high frequency. Closer inspection shows they form a regular pattern and have a $\Delta\nu$ of 6.52 d^{-1} . The strongest peak is within 1 d^{-1} of the predicted fundamental mode frequency. This ID is confirmed by the échelle diagram (Fig. 4, Section 4).

HD 103701, sector 10

In addition to the δ Sct pulsation, there are peaks at 1.35 and 1.53 d^{-1} that are likely g modes. The low p-mode frequencies are consistent with the star's position on the HR diagram, apparently beyond the TAMS (assuming solar metallicity and $f_{ov} = 0.02$). However, this star has a high RUWE (2.04) in *Gaia* DR2, so might be an unresolved binary and a little overluminous.

HD 109738, sector 11

A highly multiperiodic δ Sct star. The strongest peak is at about twice the frequency predicted by the P–L relation, making this a ‘second ridge’ star (Ziaali et al. 2019). We tentatively infer a large separation of 4.8 – 4.9 d^{-1} for this star. If the fundamental mode is excited, it is not driven strongly. A power excess around 2 d^{-1} is likely to be a group of $\ell = 1$ g modes.

HD 111786, sector 10

A known δ Sct star (MO Hya), this star exhibits pulsation over a broad range of frequencies. The strongest peak ($f_1 = 15.02 \text{ d}^{-1}$) lies close to the P–L prediction (15.37 d^{-1}). It is noteworthy that f_2 is close to (but not exactly) double the frequency of f_1 , and the two are almost equal in amplitude. That is, it would be a typical ‘second ridge’ star on the P–L diagram (Ziaali et al. 2019) if the amplitudes were only slightly different. Ridges emerge on an échelle diagram using a large frequency separation of $\Delta\nu \approx 5.2 \text{ d}^{-1}$ (Fig. A3).

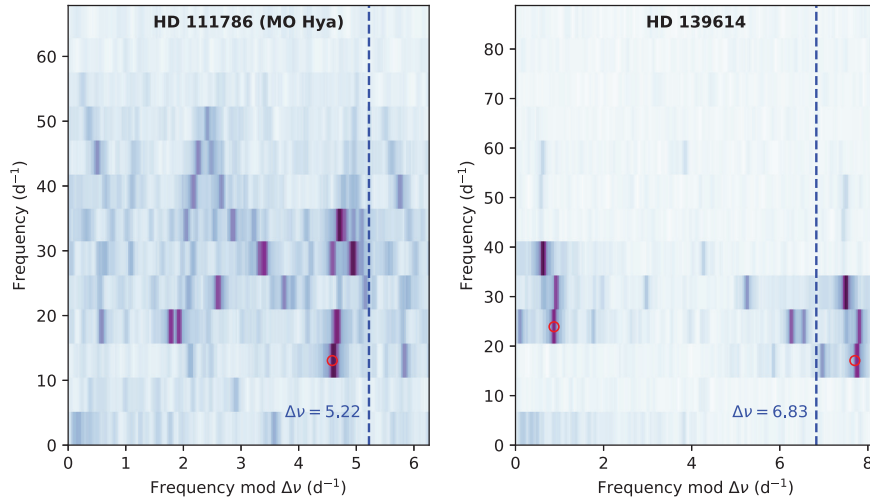


Figure A3. Échelle diagrams of HD 111786 and HD 139614. The red circles show the fundamental mode. Identification of this mode in HD 111786 is ambiguous based on the échelle only, but the P–L relation argues for the stronger of the peaks near the end of order 3.

HD 112682, sector 10

This is a non-pulsator at the 0.04 mmag level. It lies close to the theoretical blue edge of the δ Sct instability strip (Dupret et al. 2005), but comfortably inside the observational blue edge (Murphy et al. 2019).

HD 112948, sector 10

A clear δ Sct star. There is no obvious candidate for the fundamental mode. A solitary peak near 11 d^{-1} appears to be unresolved, and lies far from the prediction from the P–L relation (see Fig. 1). Two very close peaks at 1.78 d^{-1} could be g-mode pulsation, but the lack of other peaks in the typical γ Dor frequency range is puzzling if so.

HD 119896, sector 11

Only FFI data are available. This is a clear rotational variable situated close to the red edge of the δ Sct instability strip, but not a δ Sct star (at the 0.05 mmag level).

HD 120122, sector 11

A series of peaks beginning at $f = 1.49 \text{ d}^{-1}$ are near harmonics of each other, but the series is imperfect. It is unclear whether they describe rotational variability (suggesting starspots, potentially on a cooler companion) or perhaps g modes near combination frequencies (Kurtz et al. 2015; Saio et al. 2018b). We consider the unmarked peak near 4.6 d^{-1} to be part of this series. There is a Fourier peak at 16.07 d^{-1} , suggesting this is also a δ Sct star, or that there is otherwise a δ Sct star in the aperture. This peak is far from the predicted fundamental mode frequency (18.8 d^{-1}), and is low in amplitude (0.09 mmag), so we suspect it is not the fundamental mode and that this star does not pulsate in the fundamental mode. Near 8.60 d^{-1} there is a pair of poorly resolved Fourier peaks that are possibly additional pulsation modes, one of which is the strongest peak in the δ Sct frequency region. The star is located between the theoretical (Dupret et al. 2005) and observational (Murphy et al. 2019) red edges of the δ Sct instability strip.

HD 125508, sector 11

A clear δ Sct star, also with many lower frequency peaks that could be g modes, though combination frequencies of the p modes would have to be ruled out. A peak at 0.500 d^{-1} and its harmonic are present, which are not likely to be g modes. With an extremely high RUWE of 16.630 and no parallax data available in DR2, HD 125508 is almost certainly a binary or multiple system.

HD 126627, sectors 11–12

Although this is not a δ Sct star at around the 0.04 mmag level, it is a high-amplitude g-mode pulsator, with the strongest g modes achieving amplitudes of over 10 mmag. Because the time-series is only two sectors long, the g modes are not resolved from one another, but the light curve shows clear beating, and the non-sinusoidal nature of the light curve suggests the star pulsates in prograde sectoral modes and exhibits combination frequencies (Kurtz et al. 2015; Saio et al. 2018b). The pulsational properties of the star are consistent with its location on the HR diagram, at low luminosity and between the theoretical (Dupret et al. 2005) and observational (Murphy et al. 2019) red edges of the δ Sct instability strip, therefore in the middle of the γ Dor instability strip (Dupret et al. 2005).

HD 127659, sectors 11–12

The star is multiperiodic, with several peaks having amplitudes greater than 1 mmag. It was excluded from the analyses in Sections 3–5 on the basis of its hot effective temperature ($11\,900 \text{ K}$), which might suggest that it is a β Cep variable rather than a δ Sct variable. However, it has a hydrogen line (spectral) type of F2 V in Gray et al. (2017), where it was first identified as a λ Boo candidate, and has a spectral type of F2 V kA3mA4 λ Boo in Murphy et al. (submitted), who confirmed it as a λ Boo star. These suggest a temperature closer to 7000 K , in which case it is indeed a δ Sct star. It also shows low-frequency peaks in the Fourier transform, so it may be a γ Dor– δ Sct hybrid. The uncertainty over its parameters justifies its exclusion from analyses in Sections 3–5.

HD 139614, sector 12

The pulsation spectrum contains relatively few modes spread over a wide frequency range. This low mode density is amenable to analysis with only one sector of TESS data. The modes have a characteristic large spacing, which we measure to be 6.83 d^{-1} , consistent with its position near the ZAMS. The predicted fundamental mode at 21.6 d^{-1} falls very close to the observed peak at 21.36 d^{-1} . The échelle diagram (Fig. A3) confirms this is the correct identification.

HD 153747, sector 12

The TESS light curve spans only one sector, and does not have sufficient frequency resolution for the multiple pulsation modes of this star, but the three extracted frequencies (Table 1) are resolved from each other. Other peaks of similar amplitudes and frequencies are present in the Fourier transform of the light curve. We find a characteristic spacing of $\sim 4.8 \text{ d}^{-1}$, but no candidate for the fundamental mode.

HD 154153, sector 12

Not a δ Sct pulsator, down to the grass level of around 0.015 mmag . The lack of pulsation is surprising, given the star's location in the middle of the δ Sct instability strip. Two weak (0.05 mmag) low frequencies could be g modes.

HD 154951, sector 12

This δ Sct star oscillates in a single p mode (at the precision of these TESS FFI data) and apparently many g modes. The light curve appears noisy or otherwise contains considerable additional variability at the start of the sector.

HD 159021, sectors 12–13

This star seems to oscillate richly in g and p modes. The broad power excesses around the dominant p modes is not a spectral window effect. This leaves the high mode density difficult to explain. The two stronger p modes appear to be the fundamental mode ($f_2 = 10.707 \text{ d}^{-1}$) and the first radial overtone mode ($f_1 = 13.825 \text{ d}^{-1}$) based on their observed period ratio of 0.774. This suggests that the luminosity based on the *Gaia* DR2 parallax is incorrect.

HD 162193, sector 13

The light curve is of good quality. The dominant mode has a 1-h period and also shows a clear harmonic. Using the P–L relation, we predict a fundamental mode of 8.65 d^{-1} , but there is no observed peak near this frequency. The closest candidate is at 10.13 d^{-1} and there are no other contenders. The star lies between the theoretical (Dupret et al. 2005) and observational (Murphy et al. 2019) blue edges of the δ Sct instability strip.

HD 168740, sector 13

A highly multiperiodic δ Sct star. Using the P–L relation, we calculate the frequency of the fundamental mode to be 12.3 d^{-1} , which falls close to the rather prominent peak $f_2 = 12.57 \text{ d}^{-1}$. This is confirmed using an échelle diagram with a large spacing of $\Delta\nu = 4.08 \text{ d}^{-1}$ (Fig. A6). This rather small value $\Delta\nu$ is consistent with a star that has evolved considerably away from the ZAMS. Indeed,

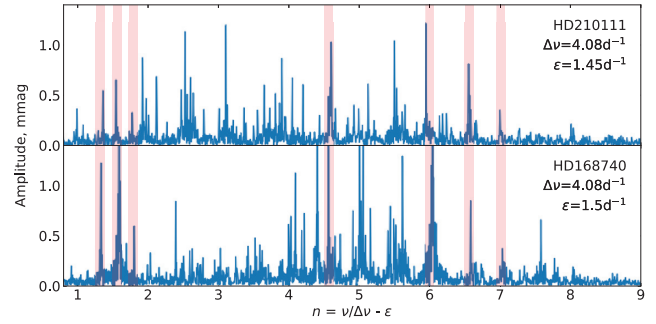


Figure A4. Amplitude spectra of the near-twins HD 210111 and HD 168740. The stars have $\Delta T_{\text{eff}} = 200 \text{ K}$ and $\Delta \log L = 0.01 \text{ dex}$. The vertical red rectangles highlight similar features within approx 0.1 d^{-1} in the two amplitude spectra. At frequencies in the middle, the pulsation spectra are too messy to make meaningful matches.

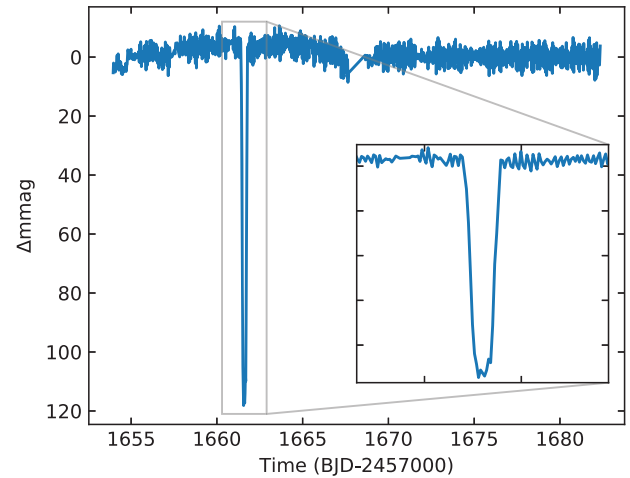


Figure A5. TESS sector 13 FFI light curve of the λ Boo star HD 168947, showing the single eclipse.

our observed T_{eff} and $\log L$ place the star about a third of the way through its main-sequence lifetime. Interestingly, a near-twin on the HR diagram (HD 210111) has a very similar amplitude spectrum (Fig. A4). A 1 mmag peak at 0.80 d^{-1} could be rotational in origin, and would be consistent with the known $v \sin i$ of 130 km s^{-1} (Heiter 2002).

HD 168947, sector 13

Only FFI data are available. The target shows a clear 0.12-mag eclipse of 0.39-d duration in the 27-d data (Fig. A5). It is difficult to determine whether the eclipse is truly flat-bottomed, but it is certainly not triangular. We infer a period of at least 21 d since only one eclipse is observed. This, and the long eclipse duration, suggest that the companion is a late-type dwarf in a wide orbit. At least one of the stars is a δ Sct variable. There are also some peaks near 0.7 d^{-1} that could potentially be r modes. Additionally, some power near 1.5 and 3 d^{-1} could be g modes, but the short light curve and low signal-to-noise make this uncertain.

HD 169142, sector 13

This target has only FFI data, the first TESS orbit of which are very poor in quality. Using only the second orbit, we infer that the star is

not a δ Sct star at the 0.15 mmag level, but it does have some low-frequency variability, probably from g modes. Although there is one significant peak above 4 d^{-1} , we did not extract it, as it is clearly associated with the lower frequency peaks rather than a p mode.

HD 169346, sector 13

This target has only FFI data. The first few days of the light curve have a significant dip, which could be the transit of a large body, or just an instrumental or data processing artefact. For the pulsation analysis we excluded these data points. The pulsation frequencies go right up to the Nyquist frequency of the FFI data, so the extracted peaks probably include aliases of super-Nyquist pulsation frequencies. The peak at 10.87 d^{-1} is close to the predicted fundamental mode frequency (10.48 d^{-1}), but it does not dominate the spectrum and the peak could be a Nyquist alias.

HD 184779, sector 13

The grouping of some p modes suggests they could be rotationally split multiplets. The star appears to have a large separation of $\Delta\nu \sim 4.0 \text{ d}^{-1}$. The dominant peaks agree with those found in SuperWASP data (Paunzen et al. 2015). We identify the strongest peak at 12.5686 d^{-1} as the fundamental mode, which for roughly solar metallicity suggests a first radial overtone mode would lie at around 16.3 d^{-1} (Petersen & Christensen-Dalsgaard 1996). However, no strong peak occurs at this frequency, leaving the mode identification as tentative. The low frequency peaks at around 0.65 and 1.90 d^{-1} could be r and g modes, respectively.

HD 188164, sector 13

Only FFI data are available, and the light curve is not particularly clean. Nonetheless, it shows the star is not a δ Sct star, at the 0.015 mmag level, despite the star sitting comfortably within the δ Sct instability strip.

HD 198160 and HD 198161, sector 13

Form a close binary with a separation of a few arcseconds (3.2 arcsec in the WDS; 4.38 arcsec in *Gaia* DR2), and a magnitude difference of 0.27 mag in *Gaia* G (Gaia Collaboration 2018). Both are known λ Boo stars (Murphy et al. 2015). These stars are unresolved by TESS so the light curve contains the pulsation properties of both stars (without guaranteeing that both are pulsators). According to *Gaia* DR2 (Gaia Collaboration 2018), both stars have the same parallax (15.3 mas) within the uncertainties. If this parallax is correct and not affected by the possible binarity of the system, then using the DR2 separation the stars are ~ 300 au apart and hence bound. Their proper motions, while similar, are not identical within the uncertainties, which likely arises from binary orbital motion. Throughout this paper, we consider only HD 198160 in our statistics, and we have 70 targets rather than 71.

HD 203709, sector 1

This is a definite δ Sct variable, and an unresolved group of peaks at $1.5\text{--}2.2 \text{ d}^{-1}$ suggests it is also a γ Dor variable. The Fourier transform reveals an unusual pulsation property, which is an almost comb-like series of peaks between 4.1 and 8.1 d^{-1} . The signal to noise of these peaks, and spacing between them at $\sim 0.7 \text{ d}^{-1} \approx 8.1 \mu\text{Hz}$, are both far too large for these to belong to a red

giant companion. The photometry is also inconsistent with one component being a red giant. There is no obvious pattern in an échelle diagram.

HD 210111, sector 1

This is a very highly multiperiodic δ Sct star, which is also discussed in Section 5. The TESS light curve has been described by Antoci et al. (2019). There seems to be a spacing of about 4.1 d^{-1} in the échelle diagram of this TESS light curve (Fig. A6), which also suggests the peak at 12.2 d^{-1} is the fundamental mode. The P–L relation agrees with this. The star is a near twin of HD 168740 (Fig. A4). The two low-frequency peaks (at 1.60 and 2.31 d^{-1}) are not harmonics of each other, which suggests one of the stars is likely a γ Dor star, too.

HD 213669, sector 1

The two strong peaks, f_2 and f_3 , near 15 d^{-1} do not have a frequency ratio consistent with them both being low-order radial modes. The predicted fundamental mode frequency lies between the two, making it difficult to determine which is the correct identification. We tentatively identify that f_2 as the stronger and lower frequency of the two peaks is the fundamental mode. We note that f_1 is not a harmonic of f_2 , but they are close. Given the frequency of f_1 , this would be a ‘second ridge’ star in the P–L diagram (Ziaali et al. 2019). It is noteworthy that the frequency content in this region of the FT is remarkably similar to HD 38043, except that the multiplet spacing is smaller in HD 38043. The low frequency peaks between 1.5 and 2.1 d^{-1} are almost certainly unresolved $\ell = 1$ g modes.

HD 214582, sector 1

An 8.7-mmag peak at 1.26 d^{-1} dominates the Fourier transform of the light curve, and is accompanied by other low frequencies of mmag amplitude. Aside from this strong low-frequency content, two p modes are evident. Both are significant (SNR = 10.3, 5.4, respectively), and only appear low in amplitude because of the contrast against the lower frequencies. This star has the highest predicted frequency from the P–L relation of all stars in the sample, but the prediction does not match the observed peaks well. We tentatively identify $f_1 = 19.76 \text{ d}^{-1}$ as the fundamental mode.

HD 223352, sector 2

The light curve contains a lot of noise, particularly of instrumental origin, but it still is possible to exclude δ Sct pulsation at the 0.02 mmag level. This is not surprising, given the T_{eff} of 9900 K, far from the δ Sct instability strip blue edge.

HD 290799, sector 6

A known δ Sct star (V1790 Ori). A clear δ Sct star with oscillations covering a wide frequency range, also discussed in Section 5 and by Bedding et al. (2020). The strongest peak ($f_1 = 22.53 \text{ d}^{-1}$) agrees with the one in SuperWASP data (Paunzen et al. 2015); both differ by 1 d^{-1} from the one found by Paunzen et al. (2002, 22.53 d^{-1}), suggesting the latter peak was a diurnal alias. Both an échelle diagram (Fig. A6) and consideration of the P–L relation suggest that this strongest peak is the fundamental mode.

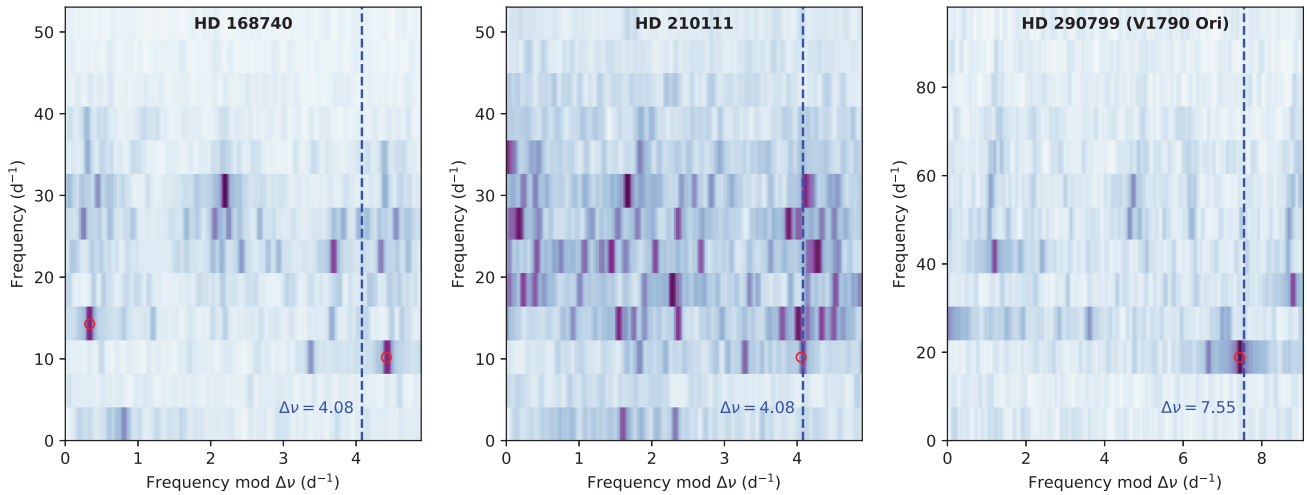


Figure A6. Échelle diagrams of HD 168740, HD 210111, and HD 290799, all of which have a clear fundamental mode (red circles).

HD 294253, sector 6

A non-variable star at the 0.05 mmag level, across all frequencies. A truly constant star, as is to be expected at this spectral type (A0). This star is also discussed in Section 5.

BD-11 1239, sector 6

This FFI-only target is highly multiperiodic with many p modes and some lower frequencies that are most likely g modes. Some of the extracted peaks could be Nyquist aliases of the real ones.

T Ori, sector 6

T Ori appears to be a δ Sct star pulsating in a single mode, at the precision of these TESS FFI data. It lies between the theoretical (Dupret et al. 2005) and observational (Murphy et al. 2019) red edges of the δ Sct instability strip. T Ori also exhibits long-period variability on time-scales of 5 d, not dissimilar in appearance to RR Lyr pulsations.

APPENDIX B: FOURIER TRANSFORMS OF λ BOO STAR LIGHT CURVES

In this online-only² appendix we show Fourier transforms of the TESS light curves. For all targets, a panel from 0 to 80 d^{-1} is shown. For stars with only FFI data, integer multiples of the Nyquist frequency are shown with vertical dashed red lines. For δ Sct stars, a zoomed plot focusing around the area of the strongest p mode is also shown adjacent to the first panel, using that same amplitude range (y-axis limit) for consistency. The extracted frequencies given in Table 1 of the main text are indicated with red diamonds, where possible. Note that peaks corresponding to the harmonics of g modes are not identified as p modes in any case.

²<https://doi.org/10.5281/zenodo.3727290>

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