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Estimation of losses in a 300 m filter cavity and quantum noise reduction in the KAGRA gravitational-wave detector

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The sensitivity of the gravitational-wave detector KAGRA, presently under construction, will be limited by quantum noise in a large fraction of its spectrum. The most promising technique to increase the detector sensitivity is the injection of squeezed states of light, where the squeezing angle is dynamically rotated by a Fabry-Pérot filter cavity. One of the main issues in the filter cavity design and realization is the optical losses due to the mirror surface imperfections. In this work we present a study of the specifications for the mirrors to be used in a 300 m filter cavity for the KAGRA detector. A prototype of the cavity will be constructed at the National Astronomical Observatory of Japan, inside the infrastructure of the former TAMA interferometer. We also discuss the potential improvement of the KAGRA sensitivity, based on a model of various realistic sources of losses and their influence on the squeezing amplitude.

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I. INTRODUCTION

Quantum noise is one of the main limiting factors for second generation gravitational-wave interferometric detectors, such as KAGRA [1], Advanced LIGO [2], and Advanced Virgo [3]. As pointed out by Caves in 1981 [4], the two manifestations of quantum noise, shot noise and radiation pressure noise, are created by vacuum fluctuations entering the interferometer by the antisymmetric (or *dark*) port. Caves also proposed the injection of a *squeezed* vacuum from the dark port as a strategy to decrease quantum noise without modifying the interferometer configuration.

In a squeezed vacuum state, the amplitude and phase uncertainty, equally distributed in an ordinary vacuum, are modified in order to reduce one at the expense of the other. If this is a pure state, it can be represented as an ellipse in the quadrature plane and it is characterized by the ratio of its axes (squeezing magnitude) and by its orientation (squeezing angle). Both of these parameters are functions of the Fourier frequency. If the quadrature with the reduced uncertainty is aligned in the gravitational-wave signal, the signal to noise ratio is improved with respect to that achievable with an ordinary vacuum. Since the optomechanical coupling of the laser light with the interferometer test masses induces a rotation of the squeezing ellipse, the

injection of a frequency-independent squeezed vacuum (i.e., with constant squeezing angle) mitigates quantum noise only in the part of the spectrum where the gravitational-wave signal is aligned with quadrature with reduced uncertainty. In the case of a phase quadrature squeezing, the improvement is the same as that achievable by increasing the laser power.

A broadband quantum noise reduction can be obtained by injecting a squeezed vacuum with an angle that varies with frequency, in such a way that the signal and reduced noise quadrature are always aligned [5]. This frequency-dependent squeezing can be obtained by reflecting off a frequency-independent squeezed state by a detuned Fabry-Pérot filter cavity. The rotation of the squeezing angle has been experimentally demonstrated in the MHz region [6] and, more recently, in the kHz region [7].

The injection of frequency-dependent squeezing is particularly suitable for a future upgrade of KAGRA, since increasing laser power to reduce shot noise can be difficult due to the cryogenic temperature of the detector. Moreover, since thermal noise is low enough to make quantum noise the dominant contribution in a large part of the spectrum, the reduction of quantum noise would have a direct and significant effect on the detector sensitivity.

The achievement of a high level of frequency-dependent squeezing faces two kinds of difficulties. First, the squeezing angle should undergo a rotation in the frequency region

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where quantum noise switches from radiation pressure noise to shot noise (~ 70 Hz for KAGRA). This corresponds to a filter cavity with a very long storage time of ~ 2 ms. Moreover, the presence of optical losses in the filter cavity system decreases the squeezing factor, since losses are associated with ordinary vacuum fluctuations that couple with the squeezed states reducing their squeezing level. For both these reasons, optical losses in the filter cavity system should be estimated and reduced as much as possible.

A 300 m prototype of the cavity will be developed at the National Astronomical Observatory of Japan (NAOJ), using the former TAMA infrastructure [8]. This is, to our knowledge, the longest filter cavity prototype under construction. Previous experiments were made using a 2 m long cavity [7]. The first goal of this paper is to present the filter cavity optical design. One of the main issues in the filter cavity realization is the losses due to the mirror surface imperfections. In this work we present a study of the specifications for the mirrors, based on real mirror maps, used in the Virgo experiment. Moreover, a crucial point is to compare the squeezing degradation from the filter cavity optical losses with that originated by other mechanisms. Thus, we present a complete *squeezing degradation budget* for the 300 m cavity, based on the work by Kwee *et al.* [9]. We finally discuss the potential improvement of KAGRA sensitivity using the 300 m filter cavity, compatible with the KAGRA infrastructure.

The interest of this work goes beyond KAGRA, since a 100 m class filter cavity is also a possible solution for a medium-term upgrade of Advanced Virgo [10], and 300 m filter cavities are planned for the third generation detector Einstein Telescope [11]. The structure of the article is as follows: in Sec. II we demonstrate the computation of the filter cavity parameters needed to obtain an optimal rotation of the squeezing angle for KAGRA. In Sec. III the filter cavity numerical simulation, with realistic mirror maps, is presented. In Sec. IV an estimation of the squeezing degradation, due to cavity losses and other degradation mechanisms, is obtained. Finally, the expected improvement in KAGRA sensitivity, achievable using frequency-dependent squeezing, is presented.

II. FILTER CAVITY OPTICAL DESIGN

In the first part of this section we discuss the choice of finesse and detuning of the filter cavity. In the second part, we address the problem of the choice of cavity length, mirror dimension, and radius of curvature (RoC).

A. Length, finesse, and detuning

Filter cavity bandwidth and detuning should be chosen in order to allow an optimal rotation angle of the squeezed state to be injected in the interferometer dark port. This rotation should counteract that which is produced by the

interferometer, keeping the squeezing state aligned with the signal quadrature at all frequencies. The optimal rotation angle of the cavity is therefore a function of the interferometer parameters [12],

$$\theta_{ic}(\Omega) = \arctan(K(\Omega)), \quad (1)$$

with K being the frequency-dependent optomechanical coupling, whose dependence from the signal sideband frequency Ω is given by

$$K(\Omega) = \left(\frac{\Omega_{\text{SQL}}}{\Omega} \right)^2 \frac{\gamma_{\text{ifo}}^2}{\Omega^2 + \gamma_{\text{ifo}}^2} \quad (2)$$

where γ_{ifo}^2 is the interferometer bandwidth and Ω_{SQL} is the frequency at which quantum noise reaches the standard quantum limit [5], marking the transition between radiation pressure noise and shot noise. In a dual recycled interferometer, with a tuned signal-recycling cavity, we have

$$\Omega_{\text{SQL}} = \left[\frac{t_{\text{sr}}}{1 + r_{\text{sr}}} \right] \frac{8}{c} \sqrt{\frac{P_{\text{arm}} \omega_0}{m T_{\text{arm}}}} \quad (3)$$

and

$$\gamma_{\text{ifo}} = \left[\frac{1 + r_{\text{sr}}}{1 - r_{\text{sr}}} \right] \frac{T_{\text{arm}} c}{4L_{\text{arm}}} \quad (4)$$

where t_{sr} and r_{sr} are the signal-recycling mirror amplitude transmissivity and reflectivity, P_{arm} is the intracavity power, T_{arm} is the arm cavity input mirror power transmissivity, ω_0 is the angular frequency of the carrier field, and m is the mirror mass.

For KAGRA, using the parameters shown in Table I, we have

$$\Omega_{\text{SQL}} = 2\pi \times 76.4 \text{ Hz} \quad (5)$$

and

$$\gamma_{\text{ifo}} = 2\pi \times 382 \text{ Hz}. \quad (6)$$

TABLE I. Values and symbols for KAGRA interferometer parameters.

Parameter	Symbol	Value
Carrier field frequency	ω_0	$2\pi \times 282$ THz
Standard quantum limit frequency	Ω_{SQL}	$2\pi \times 76.4$ Hz
Arm input mirror transmissivity	T_{arm}	0.004
Signal-recycling input transmissivity	t_{sr}^2	0.1536
Intracavity power	P_{arm}	400 kW
Mirror mass	m	22.8 kg

For $\Omega \ll \gamma_{\text{ifo}}$ Eq. (2) can be simplified and we are in the region of interest. The required rotation angle then becomes

$$\theta_{\text{fc}}(\Omega) = \arctan\left(\frac{\Omega_{\text{SQL}}}{\Omega}\right)^2. \quad (7)$$

In order to obtain the above frequency-dependent rotation, the bandwidth γ_{fc} and the detuning $\Delta\omega_{\text{fc}}$ of a lossless filter cavity should be

$$\gamma_{\text{fc}} = \frac{\Omega_{\text{SQL}}}{\sqrt{2}}, \quad (8)$$

$$\Delta\omega_{\text{fc}} = \gamma_{\text{fc}}. \quad (9)$$

As discussed in [9], in the presence of losses, the bandwidth and the detuning of the filter cavity become

$$\gamma_{\text{fc}} = \sqrt{\frac{2}{(2-\epsilon)\sqrt{1-\epsilon}}} \frac{\Omega_{\text{SQL}}}{\sqrt{2}}, \quad (10)$$

$$\Delta\omega_{\text{fc}} = \gamma_{\text{fc}}\sqrt{1-\epsilon}, \quad (11)$$

where ϵ is a function of the filter cavity round trip losses Λ_{rt}^2 , the free spectral range $f_{\text{FSR}} = c/2L_{\text{fc}}$, and Ω_{SQL} ,

$$\epsilon = \frac{4}{2 + \sqrt{2 + 2\sqrt{1 + \left(\frac{2\Omega_{\text{SQL}}}{f_{\text{FSR}}\Lambda_{\text{rt}}^2}\right)^4}}}. \quad (12)$$

Given a filter cavity with fixed length and losses, we can compute the parameter ϵ using Eq. (12). Then, using Eqs. (10)–(11), we obtain the optimal bandwidth and detuning. Finally, in order to compute the cavity finesse, we can write the bandwidth γ_{fc} of a Fabry-Pérot cavity in terms of the losses, length, and input mirror transmissivity t_{in}^2 ,

$$\gamma_{\text{fc}} = \frac{t_{\text{in}}^2 + \Lambda_{\text{rt}}^2}{2} f_{\text{FSR}}. \quad (13)$$

Inverting the previous equation, we can finally compute t_{in}^2 , which determines the cavity finesse.

The length of the filter cavity is set to 300 m, corresponding to the length of the arms of the former TAMA interferometer. The losses Λ_{rt}^2 are set to 80 ppm, a value that is justified in the following section. For these values, we found $\epsilon = 0.111$ and $\gamma_{\text{fc}} = 2\pi \times 57.3$. The optimal detuning will be $\Delta\omega_{\text{fc}} = 2\pi \times 54$ Hz and $t_{\text{in}}^2 = 0.0014$, corresponding to a finesse of 4480. We highlight the fact that for Λ_{rt}^2 up to ~ 700 ppm, the value of the finesse is almost independent of the cavity losses.

TABLE II. Filter cavity parameters.

Parameter	Symbol	Value
Length	L	300 m
Radius of curvature	RoC	variable
Mirror diameter	d	0.1 m
Input mirror transmissivity	t_{in}^2	0.0014
Finesse	F	4480
Beam diameter at waist (RoC 400 m)		0.0162 m
Beam diameter at the mirror (RoC 400 m)		0.0205 m

B. Mirror dimension and RoC

In order to reduce the losses due to the finite size of the mirrors (clipping losses), we require the size of the beam to be as small as possible. As shown in [13], this is also the best way to reduce the effects of large-scale mirror defects. The confocal configuration (radii of curvature equal to the cavity length) gives the smallest beam radius on the mirrors, but this configuration is marginally stable. We choose radii ~ 400 m, reasonably larger than 300 m, to avoid cavity instability. The exact value of the RoC is determined using numerical simulations, described in the next paragraph, in order to minimize the losses.

For a 300 m cavity with two mirrors with a RoC of 400 m, the beam diameter at the waist and on the mirrors is respectively 0.0162 m and 0.0205 m. Choosing, as it was done for TAMA, mirrors with a diameter of 0.1 m (roughly five times bigger than the beam radius) results in completely negligible clipping losses.

The filter cavity parameters are reported in Table II.

III. MIRROR SURFACE QUALITY SPECIFICATIONS

The presence of losses in the filter cavity, caused by mirror defect and mirror finite size, induces squeezing degradation. In this section we detail the procedure needed to find the specifications on mirror dimensions and surface flatness.

Mirror defects, namely deviations of the mirror surface from a perfect spherical one, are described by the mirror map, which is a square matrix of n elements with $n = l_{\text{mir}}/\text{res}$, where l_{mir} is the length of the surface described, and res is the resolution of the map. Each matrix element, corresponding to a pixel with area res^2 , contains a measure of the mirror surface height h (with respect to a perfect spherical one).

We can associate to each mirror map (or to a part of it) the root mean square (RMS) of the height, defined as

$$\sigma_{\text{RMS}} = \sqrt{\frac{1}{n} \sum_{i=1}^n (h_i - \bar{h})^2} \quad \text{with} \quad \bar{h} = \frac{1}{n} \sum_{i=1}^n h_i. \quad (14)$$

Another useful number to quantify surface flatness is the peak-to-valley value (PV): a measure of the difference between the highest and lowest point.

Mirror defects can be studied in the spatial frequency domain by applying a two-dimensional Fourier transform to the mirror map. The lowest spatial frequency f_{\min} coincides with the inverse of l_{mir} , while the maximum spatial frequency f_{\max} is given by $1/(2 \cdot \text{res})$. A one-dimensional power spectral density (PSD) can be associated with the two-dimensional Fourier transform map [14]. For such a one-dimensional PSD we have the relation

$$\sigma_{\text{RMS}}^2 = \int_{f_{\min}}^{f_{\max}} \text{PSD}(f) df. \quad (15)$$

The low frequency defects, those which have a spatial frequency up to 10^3 m^{-1} , contribute to the so-called mirror flatness while higher frequency defects are associated with the so-called mirror roughness. This distinction has no physical motivation, but it is simply due to different techniques used to measure spatial defects in the two cases [15].

The scattering angle for light at normal incidence of wavelength λ can be written as a function of the frequency of spatial defect as [16]

$$\theta = \lambda \times f. \quad (16)$$

This equation means that a defect at spatial frequency f scatters a fraction of the light at angles θ or larger. As a consequence, the amount of light reflected back at normal incidence is reduced by the same amount. This fraction is given by $(4\pi \times \sigma(f)/\lambda)^2$ where $\sigma(f)$ is the amplitude of the defect at spatial frequency f .

For a given cavity length L and mirror diameter d , there is a maximum scattering angle θ_{limit} above which light is scattered out of the cavity,

$$\theta_{\text{limit}} = \frac{d}{2L}. \quad (17)$$

Using Eqs. (16)–(17) we can then find a spatial frequency f_{limit} for the mirror defects above which the light is scattered out of the cavity,

$$f_{\text{limit}} = \frac{d}{2L \times \lambda}. \quad (18)$$

From the equation above, the losses due to defects with spatial frequency above f_{limit} can be estimated as [16]

$$\text{losses}_{(f > f_{\text{limit}})} = \left(\frac{4\pi \times \sigma}{\lambda} \right)^2 \quad (19)$$

where σ is the RMS for frequencies above f_{limit} . For the filter cavity we are considering $f_{\text{limit}} = 157 \text{ m}^{-1}$.

It is important to note that the RMS for frequencies lower than f_{limit} also contributes to losses. In fact, even if light is not immediately scattered out of the cavity, it is likely to be transferred on higher order modes and eventually exit the cavity. An accurate estimation of the round trip losses has been done using a numerical simulation of the fields in the cavity after adding realistic maps to the mirror surfaces. This allowed us to set specifications for the mirror quality needed to keep losses below a desired threshold.

The round trip losses in a Fabry-Pérot cavity are defined as [17]

$$\Lambda_{\text{rt}}^2 = \frac{P_{\text{in}} - P_r - P_t}{P_{\text{circ}}} \quad (20)$$

where P_{in} is the input power (which is assumed to be a fundamental mode), and P_{circ} , P_t , P_r are the powers circulating in the cavity, transmitted and reflected, respectively. Since we can only take advantage of the light reflected on the fundamental mode, the definition of Eq. (20) has been modified to

$$\Lambda_{\text{rt}}^2 = \frac{P_{\text{in}} - P_r^{00}}{P_{\text{circ}}} \quad (21)$$

where P_r^{00} is the fraction of the reflected power that is on the fundamental mode.

We used the MATLAB package OSCAR [18] to simulate the fields and compute the values of P_{circ} and P_r^{00} to be used in Eq. (21).

We ran the simulation using five different mirror maps, for mirrors used in the Virgo experiment. Four of these mirrors were produced for the initial Virgo experiment (2007–2011) with a standard polishing technology. The fifth belongs to the Advanced Virgo experiment (which will become operational in 2016) and is obtained with an ion beam polishing technique. The Virgo maps have a resolution of about $350 \mu\text{m}$, while the Advanced Virgo map has a resolution of $378.4 \mu\text{m}$. One example of the Virgo maps and the Advanced Virgo map are shown in Fig. 1 along with their relative PSD.

The cavity parameters used in the simulation are those reported in Table II. The maps were only applied to the end mirror, while the input mirror has been considered perfect. We checked that the round trip losses for a cavity where both the mirrors have defects can be obtained by simply multiplying the previous result by 2. Results presented here have already been multiplied by 2.

Each surface has been characterized by its RMS and its PV over different diameters. The measured values are reported in Table III. The four mirror of initial Virgo are referred to as V1, V2, V3, V4, while the mirror of Advanced Virgo is called ADV. Round trip losses for the various mirrors have been calculated in terms of the radius of curvature and are reported in Fig. 2. The loss floor

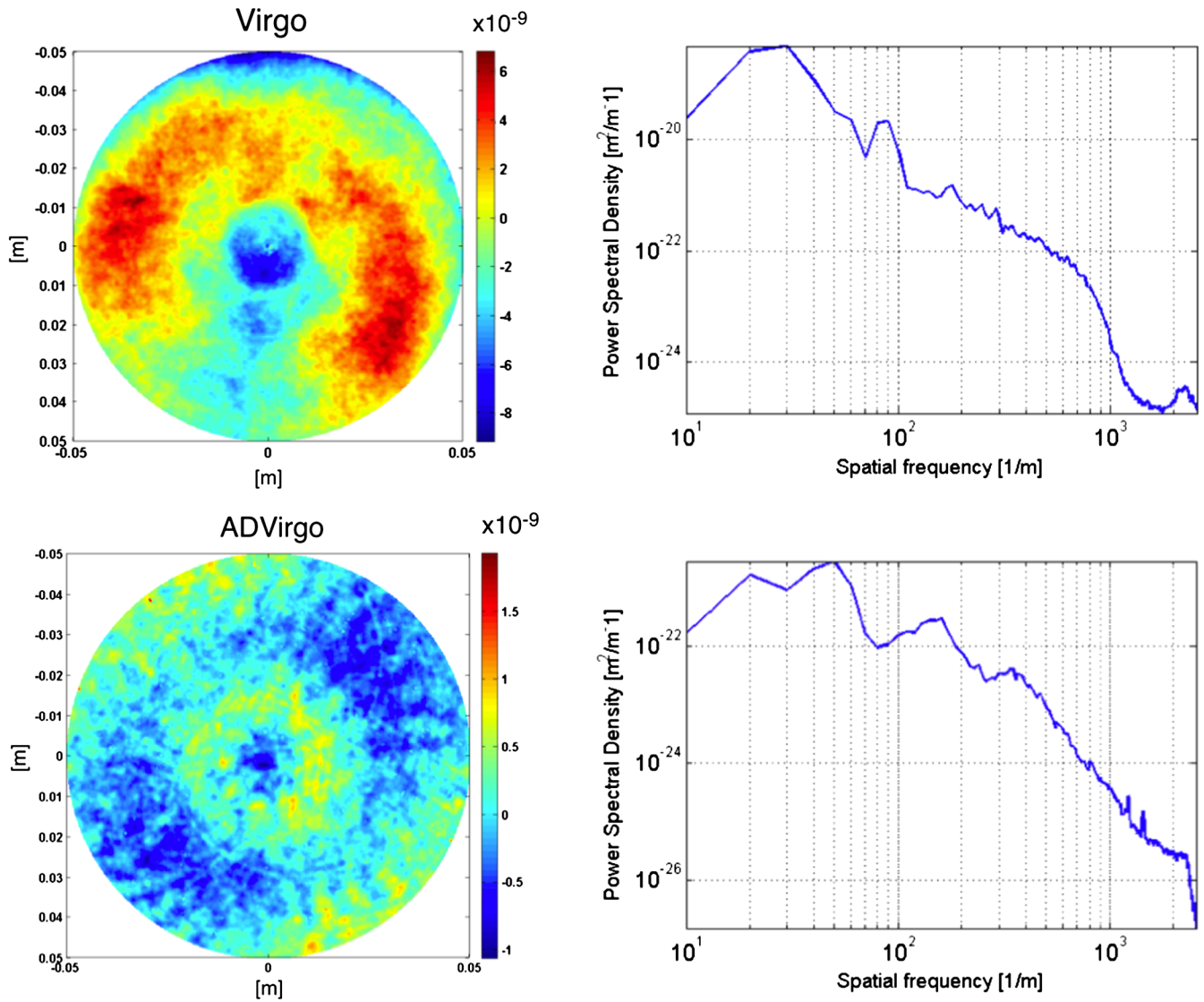


FIG. 1. Initial Virgo map (top) and Advanced Virgo map (bottom), along with mirror maps (left) and PSD (right).

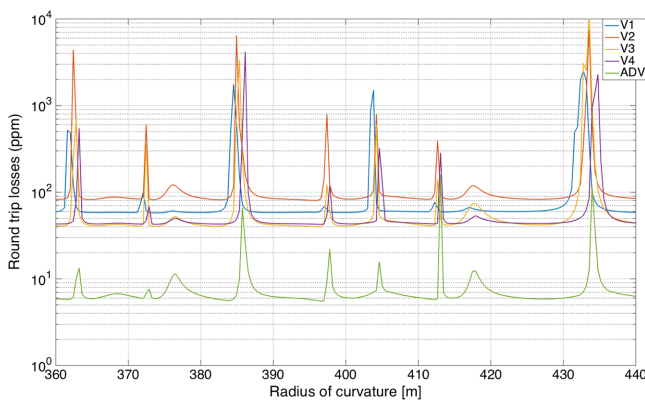


FIG. 2. Round trip losses as a function of the radius of curvature. Peaks correspond to values of RoC for which the cavity is quasigenerate.

for each mirror has also been reported in the last column of Table III in order to be directly compared with the mirror flatness. The presence of peaks in the plots of Fig. 2 is due to power transferred to higher order modes that are partially resonant along with the fundamental mode for certain values of the curvature radius.

Maps used in the simulation account for mirror defects with spatial frequency going from 10 m⁻¹ to 2 × 10³ m⁻¹. This means that losses caused by mirror roughness are not included in this estimation. A map of the roughness has been measured for the Advanced Virgo mirror. This map, obtained by scanning an area of 0.3 mm × 0.3 mm with a resolution of 1.28 μm, scans frequencies from 3.3 × 10³ to 3.9 × 10⁵ m⁻¹. From its RMS, under the assumption that it is uniform on the surface, we can estimate additional losses due to roughness, which should be added to those already measured. The RMS is 0.08 nm, which corresponds to

TABLE III. RMS and PV (over different diameters) and the round trip loss floor for each mirror map. The values indicated for the losses correspond to the floor of Fig. 4.

Mirror	Diameter 0.10 m		Diameter 0.05 m		Diameter 0.02 m		Diameter 0.01 m		Losses (ppm)
	RMS (nm)	PV (nm)	RMS (nm)	PV (nm)	RMS (nm)	PV (nm)	RMS (nm)	PV (nm)	
V1	2.617	15.95	1.424	9.46	0.687	6.04	0.558	5.38	57.8
V2	1.875	15.64	1.234	8.56	0.682	6.29	0.812	5.92	80.6
V3	2.499	15.34	1.360	10.51	0.754	4.31	0.430	3.31	39.8
V4	1.752	45.61	0.984	12.12	0.509	4.46	0.531	4.46	42.6
ADV	0.319	2.99	0.274	2.09	0.192	1.18	0.142	0.97	5.6

0.89 ppm of additional losses for a single reflection. For the initial Virgo mirrors, the specification on the roughness RMS was 0.1 nm, which corresponds to 1.4 ppm of additional losses for a single reflection. In both cases the losses are dominated by flatness defects. The roughness map for Advanced Virgo and the relative PSD are shown in Fig. 3. The measurement of the light scattered at angles larger than a few degrees gives losses of the order of 5 ppm. These include both the losses due to the roughness discussed above and those due to point defects. The corresponding additional round trip losses are about 10 ppm. We remark that, even including all these effects, a difference still exists between the measured losses in Advanced LIGO and simulation results. Losses due to scattering at angles between mrad and a few degrees are being investigated as a possible cause of these differences [19]. These hypothetical losses are not included in this paper.

The conclusions of this study are, for this filter cavity (length 300 m, mirror diameter 10 cm, and RoC ~ 400 m), the following:

- (i) An Advanced Virgo-class mirror will produce floor losses < 10 ppm and a Virgo-class mirror will produce losses $\sim 40 - 80$ ppm. In order to determine

the final specifications on the mirror flatness, the squeezing degradation given by the cavity losses has to be compared to the other degradation mechanisms. This analysis is performed in the next section.

- (ii) The accidental degeneracy can amplify the losses by more than an order of magnitude. The simulation gives the *safe* regions, where the losses are at the floor level.
- (iii) The precise RoC value should be chosen in some of the floor loss regions. A precision on the RoC of $\sim 1\%$ is necessary to guarantee the RoC to be in these regions.

A. Mirror dimensions

Choosing the best mirror dimension is not as straightforward as it may seem to be. The scattering mechanisms described in the previous paragraph show that bigger mirrors have lower levels of losses. In fact, we see from Eq. (18) that by increasing the mirror diameter d , the minimum spatial frequency f_{limit} of defects which scatter light out of the cavity is higher. Consequently, the RMS for frequencies above f_{limit} in Eq. (19) is reduced and so too are the losses. However, this is not the only effect to be taken

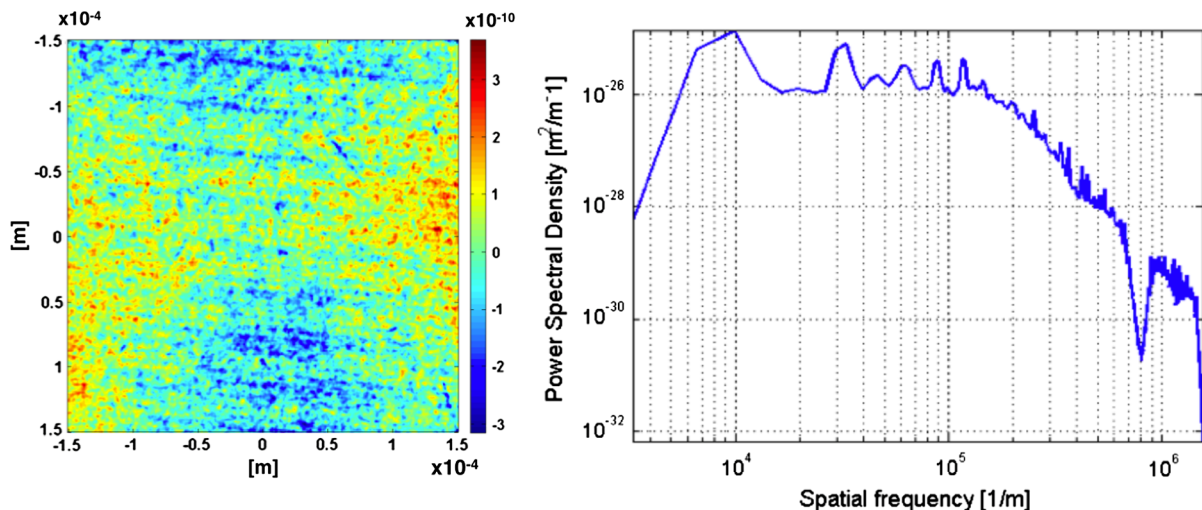


FIG. 3. Roughness maps for the Advanced Virgo configuration, mirror map (left), and PSD (right).

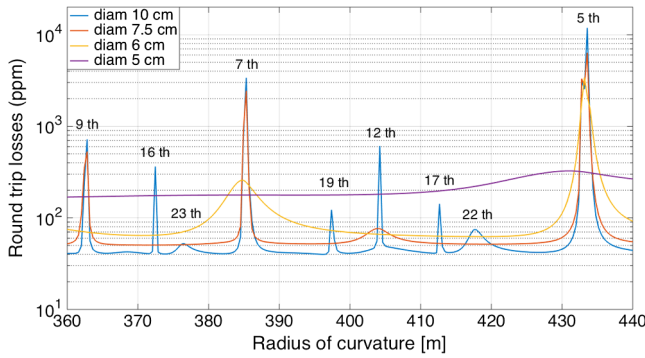


FIG. 4. Round trip losses as a function of RoC for different values of the mirror diameter. Peaks that correspond to cavity degeneracies are reduced for smaller mirrors while the floor of round trip losses increases.

into account. We have already observed that for certain values of the curvature radius, the cavity can degenerate. This means that the separation between the resonance frequency of the fundamental mode and that of a higher order mode is small enough to make it partially resonant. Some power of the fundamental mode is then transferred to this higher order mode, and losses in the fundamental mode are remarkably increased. Critical RoC values are highlighted by simulation and correspond to the peak observed in Fig. 2. Mirror dimensions also have a strong impact on the appearance of such peaks. In Fig. 4, round trip losses for a Virgo mirror (V3) are shown as a function of the RoC for different values of the mirror diameter. We see that for bigger diameters more peaks are observed, i.e., there are more RoC values that make the cavity degenerate. This number is reduced for smaller mirrors. This effect can be possibly explained by considering the intensity profile of higher order modes. In fact, the power of the higher mode is spread on a bigger surface than the fundamental mode. Generally, the *width* of a mode increases with its order, and then a reduction of the mirror dimension prevents wider modes from resonating. This explanation is confirmed by the fact that a gradual reduction of mirror dimension first eliminates resonances of the modes with higher n . Using a smaller mirror reduces the number of peaks, while increasing the floor loss level according to the mechanism explained at the beginning of the paragraph. Therefore, mirror dimensions should be chosen in order to strike the best balance between low floor losses and the presence of a safe zone for the RoC reasonably far from degeneracies. Using smaller mirrors allowed us to relax accuracy requirements on the RoC value. We can see from Fig. 4 that a mirror diameter of 0.075 m allows us to pick a RoC, for example, of 420, for which losses remain basically constant in a range of ~ 20 m. Therefore, the required precision on the RoC is $\sim 5\%$. For such a mirror dimension the contribution of clipping losses to cavity losses is still negligible. An initial value for the mirror dimension has been set to 0.1 m. This was the diameter of the former TAMA mirrors and it allows

the reuse of part of the suspension system already developed for TAMA. If needed, diaphragms can be placed in front of the mirror to reduce their diameter according to what we have observed.

B. Simple round trip simulation

A complete simulation of the cavity has been used to compute fields in our system and then to evaluate round trip losses, according to Eq. (21). This simulation allows us to take into account the effect of higher order mode resonance. In this case, light does not exit the cavity but is partially transferred to a higher order mode. This effect is observable by comparing the round trip losses on the fundamental mode with those in all modes (i.e., the total amount of light exiting the cavity). Coinciding with resonance, the curves relative to the two cases show a discrepancy accounting for the light that is still in the cavity but not on the fundamental mode. On the other hand, when the cavity is not degenerate, the total amount of light exiting the cavity coincides with the light lost on the fundamental mode. This quantity corresponds to the floor loss level reported in Table IV and has been compared with that obtained by simply computing the power escaping the mirror aperture after a single round trip [reflection from a mirror with real map, propagation for 300 m, then reflection again from a perfect mirror (as in the simulated cavity) and propagating again for 300 m]. In practice, we measure the lost power after a cavity round trip of the beam. The comparison between losses obtained with the complete simulation and those found with the simple round trip is shown in Table IV for different mirrors maps and for different values of the diameter and shows a good agreement. Even if the second method is less accurate and cannot be used in the case of degeneracy, it is much faster and can be employed to obtain a rough estimation of the floor level of the round trip losses in the cavity. A plot of the light power exiting cavity after the first reflection is shown in Fig. 5.

TABLE IV. Comparison between round trip losses (RTL) obtained with the complete simulation and those found with simple round method for different mirrors maps and for different value of the diameter. Results of the quick simulation have been multiplied by two as in the previous case.

Mirror	Diameter	RTL—full simulation	RTL—quick simulation
V1	10 cm	57.8 ppm	57.6 ppm
V2	10 cm	80.6 ppm	79.4 ppm
V4	10 cm	42.6 ppm	40.4 ppm
ADV	10 cm	5.6 ppm	5.6 ppm
V3	10 cm	39.8 ppm	38.8 ppm
V3	7.5 cm	50.4 ppm	50.0 ppm
V3	6.0 cm	62.4 ppm	62.0 ppm
V3	5.0 cm	168.6 ppm	165.4 ppm

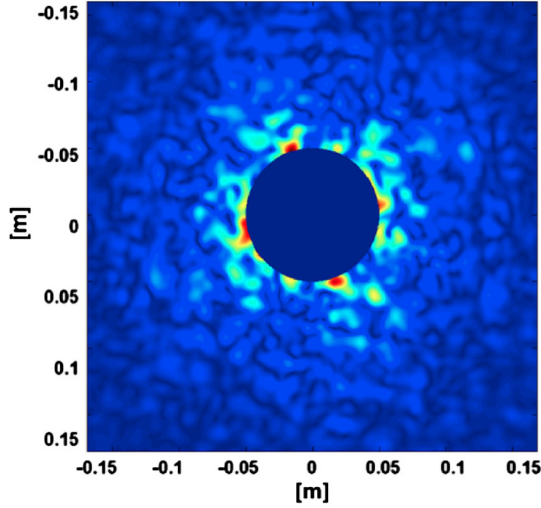


FIG. 5. Amplitude of the field scattered out of the cavity by high spatial frequency defects. The field is observed after being reflected from a mirror with a real map and then propagated for 300 m.

IV. SQUEEZING DEGRADATION AND DECOHERENCE ESTIMATION

The quantum noise in an interferometer, normalized with respect to shot noise, can be simply expressed as

$$N(\Omega) = 1 + K^2(\Omega) \quad (22)$$

where K is the optomechanical coupling defined in Eq. (2).

In an ideal system, the injection of a frequency-dependent squeezed vacuum from the interferometer dark port reduces the quantum noise to

$$N(\Omega) = e^{-2\sigma}(1 + K^2(\Omega)) \quad (23)$$

where σ is connected with the squeezing magnitude usually expressed in decibel by $\sigma_{\text{dB}} = -20\sigma \times \log_{10} e$.

In a real setup, two factors prevent this optimal noise reduction: first, optical losses decrease the squeezing level, introducing a nonsqueezed vacuum; second, fluctuations of the squeezing angle preclude an optimal rotation of the squeezed state.

In [9] a detailed analysis of several of these mechanisms has been applied to a 16 m filter cavity, with round trip losses of 1 ppm/m, which is considered a possible short-term solution for Advanced LIGO [20]. The same analysis is performed here for the 300 m filter cavity, with round trip losses of 80 ppm, corresponding to a conservative estimate for Virgo-class quality mirrors. To ease the comparisons between the two cases, the numerical values for other sources of squeezing degradation are exactly the same of [9], as reported in Table V.

Figure 6 represents the squeezing degradation budget for the 300 m filter cavity. It shows the ratio between the

TABLE V. Parameters used in the estimation of squeezing degradation.

Parameter	Symbol	Value
Filter cavity losses	Λ_{ft}^2	80 ppm
Injection losses	Λ_{inj}^2	5%
Readout losses	Λ_{ro}^2	5%
Mode-mismatch squeezer-filter cavity	Λ_{mmFC}^2	2%
Mode-mismatch squeezer-local oscillator	Λ_{mmLO}^2	5%
Filter cavity length noise (RMS)	δL_{fc}	0.3 pm
Injected squeezing	σ_{dB}^2	9 dB

quantum noise of a dual recycled Fabry-Pérot Michelson interferometer without squeezing (for instance, initial KAGRA) and the quantum noise in the presence of frequency-dependent-squeezing. An initial realistic squeezing level of 9 dB has been considered and the various degradation mechanisms have been taken into account separately and combined (black curve). In this analysis, as for [9], the contribution of the interferometer losses has been neglected. The squeezing degradation mechanisms considered are explained as follows (for a detailed description see [9]).

A. Filter cavity losses

The optical losses in the filter cavity spoil the squeezing in two ways: first, they corrupt squeezing with antisqueezing, mixing the two squeezing quadratures. Secondly, they mix standard vacuum with squeezed vacuum. For a 16 m cavity, the filter cavity losses represent the main contribution to squeezing degradation up to 100–200 Hz. Since this effect depends on the round trip losses per meter [21], it is considerably reduced in a 300 m cavity with round trip losses of 80 ppm, and at low frequencies becomes comparable with that of the mode mismatching. For this reason it is not useful to further reduce filter cavity losses

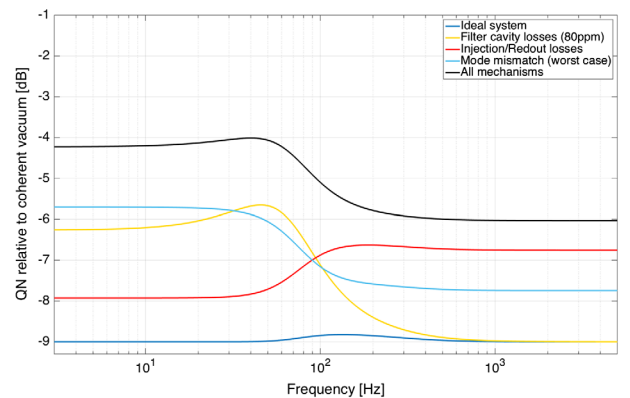


FIG. 6. Squeezing degradation budget. Quantum noise relative to coherent vacuum in the signal quadrature for an ideal system (blue curve) is compared with that obtained taking into account degradation mechanisms (one by one).

by increasing the mirror quality (for example, using Advanced Virgo-class mirrors), unless mode matching is substantially improved.

B. Injection and readout losses

Injection losses, Λ_{inj}^2 (caused by scattering, absorption, and imperfections in the optics), and readout losses, Λ_{ro}^2 (from the interferometer to the readout, including the photodetector quantum efficiency), cause a squeezing degradation by mixing ordinary vacuum with squeezed vacuum. Being independent of the cavity length, their impact does not change with a longer cavity. This mechanism is the dominating source above 100–200 Hz, assuming a loss value of 5% both for injection and readout losses. In this region, quantum noise is by far the limiting noise: a reduction in injection/readout losses leads to a consistent improvement of the detector sensitivity.

C. Mode mismatching

Following the analysis shown in [9], a squeezing degradation is also determined by an imperfect mode matching between the squeezed field and the cavity mode (Λ_{mmFC}^2), and between the cavity mode and the local oscillator (Λ_{mmLO}^2). This mismatching allows part of the field to bypass the cavity without experiencing frequency rotation, and is also a source of losses.

The magnitude of the mode mismatching can be easily measured. Nevertheless, the amplitude of the squeezing degradation depends of an arbitrary phase that is difficult to quantify. For this reason, in Fig. 2 we have shown the worst case scenario.

As in the case of injection and readout losses, this effect does not depend on the filter cavity length and, for a 300 m filter cavity with a RTL of about 80 ppm, it is comparable with the degradation due to filter cavity losses. The estimation is done assuming a mismatch of 2% between the squeezed injected field and filter cavity modes and a mismatch of 5% between the injected field and local oscillator.

D. Frequency-dependent phase noise

A length noise of the filter cavity (possibly caused by seismic noise or control noise) δL results in a shift of the optimal detuning $\delta\Delta\omega_{\text{fc}}$ given by

$$\delta\Delta\omega_{\text{fc}} = \frac{\omega_0}{L} \delta L. \quad (24)$$

This compromises an optimal squeezing quadrature rotation, limiting the achievable noise reduction. In a 300 m cavity, assuming a δL (RMS) of about 0.3 pm, this effect becomes completely negligible with respect to other mechanisms.

V. IMPROVEMENT IN KAGRA SENSITIVITY

This estimation of the achievable level of frequency-dependent squeezing using a 300 m filter cavity can be used to quantify the related improvement in KAGRA sensitivity. In Fig. 7 the quantum noise for KAGRA, without squeezing, is compared with the quantum noise obtained using 9 dB of frequency-dependent squeezed light, when all degradation mechanisms previously described are taken into account. We considered both the case of a filter cavity with round trip losses of 80 ppm and that of a perfect filter cavity. The comparison shows that no major improvements can be obtained by reducing round trip losses under the level of ~ 80 ppm since their effect becomes comparable with that of other degradation mechanisms.

Figure 8 shows the improvement in KAGRA sensitivity using 9 dB frequency-dependent squeezing, considering a

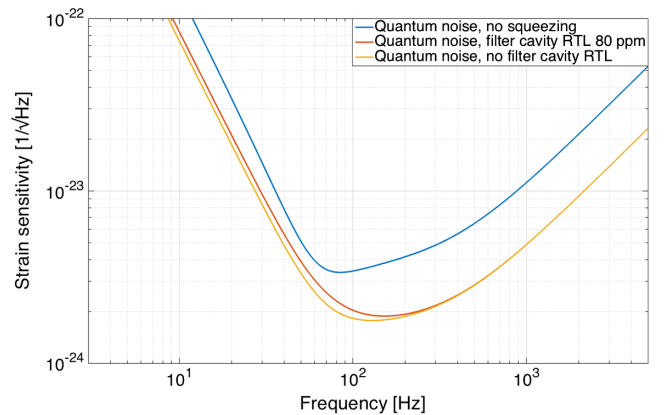


FIG. 7. Quantum noise for KAGRA without squeezing, compared with the quantum noise using 9 dB frequency-dependent squeezing, both in the ideal system and when all the degradation mechanisms are taken into account. We remark that quantum noise without squeezing is relative to official KAGRA sensitivity using a configuration with a homodyne detection angle of 121.8° , while quantum noise in the presence of squeezing uses a standard homodyne angle of 90° .

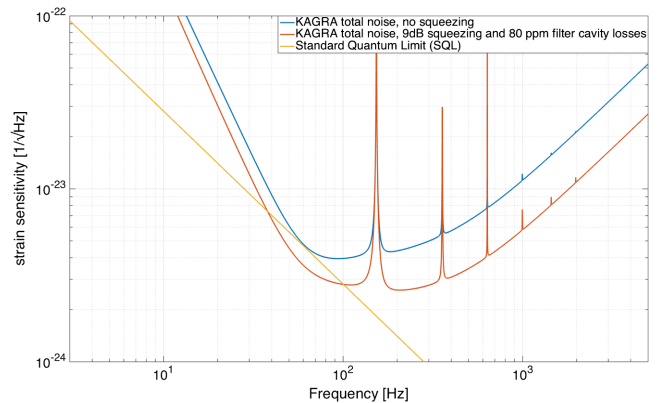


FIG. 8. Improvement in KAGRA sensitivity using 9 dB frequency-dependent squeezing, considering lossy cavity and other degradation mechanisms.

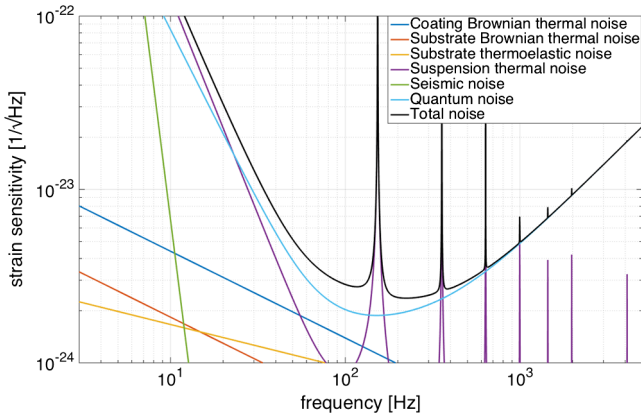


FIG. 9. The quantum noise in the presence of 9 dB squeezing and lossy system (filter cavity RTL 80 ppm) is compared with the other noise sources. Note that below 100 Hz, the contribution from thermal noise would prevent an improvement in the total sensitivity even in the case of a further reduction of quantum noise.

filter cavity with round trip losses of 80 ppm and other degradation mechanisms. We remark that the use of squeezing allows us to reach a sensitivity beyond the standard quantum limit around 70 Hz. In Fig. 9 the quantum noise of a realistic lossy system is shown along with other KAGRA noise sources. We observe how a reduction in only the quantum noise results in an improvement in almost the whole KAGRA observation bandwidth.

We also make the following remarks:

- (i) In the frequency region below 100 Hz, thermal noise is close to quantum noise and there is little to be gained from a further significant reduction in quantum noise (obtainable by improving mismatching and injection/readout losses, and decreasing losses in the filter cavity).
- (ii) In the frequency region above 100 Hz, which is dominated by quantum noise, there is much more room for improvements. However, a reduction of the filter cavity optical losses would not provide a higher squeezing level, since in this region the squeezing degradation is mainly caused by injection and readout losses.

This suggests that a further reduction of the filter cavity losses would not be necessary, while a major benefit can be obtained by protecting squeezing from losses due to injection, detection, and mismatching.

VI. SUMMARY

We have presented the optical design of a 300 m filter cavity used for the KAGRA experiment. The cavity will be tested at the National Astronomical Observatory of Japan, inside the TAMA infrastructure. First, given KAGRA parameters, we have found the detuning and finesse values needed to produce an optimal rotation of the squeezing angle. In order to find the surface quality needed for the filter cavity, we have computed the round trip losses using a numerical simulation with realistic mirror maps. The test bench maps used are those of Virgo and Advanced Virgo.

An analysis of the squeezing deterioration mechanism shows that for a round trip loss value of 80 ppm, corresponding to a conservative estimate for Virgo-class mirrors, the squeezing degradation due to the filter cavity is comparable or lower than contributions from the cavity mismatching and injection/detection losses. Moreover, frequency-dependent phase noise becomes completely negligible.

We then computed the improvements in KAGRA sensitivity with a 9 dB frequency-dependent squeezing and realistic loss mechanisms. We found a sensitivity improvement in almost the full KAGRA signal bandwidth, with a factor ~ 2 at high frequency. As already mentioned, the interferometer internal losses have been neglected in this analysis. In order to have a more precise estimation of the sensitivity improvement achievable using frequency-dependent squeezing, the contribution of these losses should be estimated.

We have also shown that, without an improvement in other noise sources, a further reduction of the filter cavity losses does not lead to a considerable improvement in the detector sensitivity.

The results of this work can be easily extended to any gravitational-wave interferometric detector planning the use of a 100 m class filter cavity, and in particular for a medium term upgrade of Advanced Virgo and for Einstein Telescope.

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- [1] Y. Aso, Y. Michimura, K. Somiya, M. Ando, O. Miyakawa, T. Sekiguchi, D. Tatsumi, and H. Yamamoto (KAGRA Collaboration), Interferometer design of the KAGRA gravitational wave detector, *Phys. Rev. D* **88**, 043007 (2013).
- [2] J. Aasi *et al.*, The LIGO Scientific Collaboration, Advanced LIGO, *Classical Quantum Gravity* **32**, 074001 (2015).
- [3] F. Acernese *et al.*, The Virgo Collaboration, Advanced Virgo: A second generation interferometric gravitational-wave detector, *Classical Quantum Gravity* **32**, 024001 (2015).
- [4] C. M. Caves, Quantum-mechanical noise in an interferometer, *Phys. Rev. D* **23**, 1693 (1981).
- [5] H. J. Kimble, Y. Levin, A. B. Matsko, K. S. Thorne, and S. P. Vyatchanin, Conversion of conventional gravitational-wave interferometers into quantum nondemolition interferometers by modifying their input and/or output optics, *Phys. Rev. D* **65**, 022002 (2001).
- [6] S. Chelkowski, H. Vahlbruch, B. Hage, A. Franzen, N. Lastzka, K. Danzmann, and R. Schnabel, Experimental characterization of frequency-dependent squeezed light, *Phys. Rev. A* **71**, 013806 (2005).
- [7] E. Oelker, T. Isogai, J. Miller, M. Tse, L. Barsotti, N. Mavalvala, and M. Evans, Audioband Frequency-Dependent Squeezing, *Phys. Rev. Lett.* **116**, 041102 (2016).
- [8] M. Ando *et al.*, Stable Operation of a 300-m Laser Interferometer with Sufficient Sensitivity to Detect Gravitational-Wave Events within Our Galaxy, *Phys. Rev. Lett.* **86**, 3950 (2001).
- [9] P. Kwee, J. Miller, T. Isogai, L. Barsotti, and M. Evans, Decoherence and degradation of squeezed states in quantum filter cavities, *Phys. Rev. D* **90**, 062006 (2014).
- [10] Advanced Virgo Squeezing Working Group, Advanced Virgo squeezer technical design report, Virgo Report No. VIR-0411A-15, 2015.
- [11] M. Abernathy *et al.*, Einstein gravitational-wave telescope conceptual design study, <http://www.et-gw.eu/etdsdocument>.
- [12] J. Harms, Y. Chen, S. Chelkowski, A. Franzen, H. Vahlbruch, K. Danzmann, and R. Schnabel, Squeezed-input, optical-spring, signal-recycled gravitational-wave detectors, *Phys. Rev. D* **68**, 042001 (2003).
- [13] F. Magaña-Sandoval, R. X. Adhikari, V. Frolov, J. Harms, J. Lee, S. Sankar, P. R. Saulson, and J. R. Smith, Large-angle scattered light measurements for quantum-noise filter cavity design studies, *J. Opt. Soc. Am. A* **29**, 1722 (2012).
- [14] H. Yamamoto, One-dimensional PSD of mirror maps, LIGO Report No. LIGO-T1100353-v1, 2011.
- [15] N. Straniero, J. Degallaix, R. Flaminio, L. Pinard, and G. Cagnoli, Realistic loss estimation due the mirror surfaces in a 10 meters-long high finesse Fabry-Pérot filter cavity, *Opt. Express* **23**, 21455 (2015).
- [16] J. C. Stover, *Optical scattering: Measurement and analysis* (SPIE Press, 1995).
- [17] J. Degallaix *et al.*, Defining the arm cavity loss for Advanced Virgo, Virgo Report No. VIR-706A-10.
- [18] J. Degallaix, OSCAR a Matlab based optical FFT code, *J. Phys. Conf. Ser.* **228**, 012021 (2010).
- [19] H. Yamamoto (private communication).
- [20] M. Evans, L. Barsotti, P. Kwee, J. Harms, and H. Miao, Realistic filter cavities for advanced gravitational-wave detectors, *Phys. Rev. D* **88**, 022002 (2013).
- [21] F. Ya. Khalili, Optimal configurations of filter cavity in future gravitational-wave detectors, *Phys. Rev. D* **81**, 122002 (2010).