

Chemical abundances along the 1G sequence of the chromosome maps: The Globular Cluster NGC 3201*

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ABSTRACT

The *Hubble Space Telescope* (*HST*) UV Legacy Survey of Galactic Globular Clusters (GCs) has investigated multiple stellar populations by means of the “chromosome map” (ChM) diagnostic tool that maximises the separation between stars with different chemical composition. One of the most challenging features revealed by ChMs analysis is the apparent inhomogeneity among stars belonging to the first population, a phenomenon largely attributed to He variations. However, this explanation is not supported by the uniformity in *p*-capture elements of these stars. The *HST* survey has revealed that the GC NGC 3201 shows an exceptionally wide coverage in the $\Delta_{F275W,F814W}$ parameter of the ChM. We present a chemical abundance analysis of 24 elements in 18 giants belonging to the first population of this GC, and having a wide range in $\Delta_{F275W,F814W}$. As far as the *p*-capture elements are concerned, the chemical abundances are typical of 1G stars, as expected from the location of our targets in the ChM. Based on radial velocities and chemical abundances arguments, we find that the three stars with the lowest $\Delta_{F275W,F814W}$ values are binary candidates. This suggests that, at least those stars could be explained with binarity. These results are consistent with evidence inferred from multi-band photometry that evolved blue stragglers populate the bluest part of the 1G sequence in the ChM. The remaining 15 spectroscopic targets show a small range in the overall metallicity by ~ 0.10 dex, with stars at higher $\Delta_{F275W,F814W}$ values having higher absolute abundances. We suggest that a small variation in metals and binarity govern the color spread of the 1G in the ChM, and that evolved blue stragglers contribute to the bluest tail of the 1G sequence.

Keywords: globular clusters: individual (NGC 3201) — chemical abundances – Population II – Hertzsprung-Russell diagram

1. INTRODUCTION

The presence of multiple stellar populations in globular clusters (GCs) is a well-assessed fact. Chemical abundance variations within stars in a given GC have been known for many years now, with some of the earliest studies including Popper (1947), Harding (1962), Osborn (1971), Cohen (1978). More recently, different stellar populations have been detected through multiple sequences along the color-magnitude diagram (CMD, e.g. Milone et al. 2012a). The photometrically-observed multiple sequences are associated with variations in chemi-

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cal elements involved in hot-H burning (light elements, CNO_{Na}, see [Marino et al. 2008](#); [Yong et al. 2008](#), for the earliest studies on this issue). One stellar population, usually considered the first population (or generation), is chemically similar to field halo stars, while the other populations display various degrees of enrichment in He/N/Na, depletion in C, and depletion of O with respect to the first population, as opposite to the α -enhancement typical in halo field stars (e.g. [Carretta et al. 2009](#); [Gratton et al. 2012](#)).

On top of this typical observed pattern, the recent analysis of UV high-precision data from the *Hubble Space Telescope* (*HST*) has revealed an even more complex picture ([Milone et al. 2015, 2017](#)). All the GCs exhibit two main discrete groups of first-generation (1G) and second-generation (2G) stars along the “chromosome map” (ChMs), that constitute the most effective and successful diagnostic tool to isolate the different populations of stars hosted in a GC. These sort of two-colour photometric diagrams, constructed by combining multi-filter *HST* images in m_{F275W} , m_{F336W} , m_{F438W} , m_{F814W} , are highly sensitive to the chemistry of the different stellar populations. On a typical ChM plane, $\Delta_{C\ F275W,F336W,F438W}$ vs. $\Delta_{F275W,F814W}$ ¹, all the observed GCs display a typical ChM shape, with a different number of seemingly-discrete groups, appearing along the main pattern, that are associated with distinct stellar populations. By combining spectroscopy and photometry we find that 1G stars on the ChMs, with low $\Delta_{C\ F275W,F336W,F438W}$ values, share the same chemical composition of halo field stars at the same metallicity, while 2G stars, are enhanced in He, N and Na and depleted in C and O ([Marino et al. 2019](#)).

The ChMs exhibit a variety of morphologies both in terms of number of stellar populations and in terms of shape and extension. [Milone et al. \(2017\)](#) subdivided 57 analysed maps in two main groups: (i) Type I GCs having a single ChM pattern, with a 1G and two or more 2G populations, and (ii) Type II GCs displaying multiple ChMs, with minor populations located on red additional ChMs. While we refer to [Marino et al. \(2019\)](#) for a detailed chemical characterisation of these two main classes of maps, here we just recall that the additional red stellar populations in Type II GCs are enhanced in metallicity, and (in some cases) in the elements produced via *slow*-neutron capture reactions (*s*-elements, e.g. [Yong & Grundahl 2008](#); [Marino et al. 2009, 2015](#)).

ChMs, besides being a powerful tool to explore the multiple stellar population phenomenon, provide fundamental information that can shed light on how GCs formed and evolved, which might be hidden in the 1G stars themselves. Indeed, one of the most interesting features observed on the ChMs is perhaps the apparent chemical inhomogeneity within the 1G population. Indeed the 1G sequence on the ChM is either elongated or bimodal, thus indicating that its stars are not consistent with a simple stellar population. [Milone et al. \(2015, 2018\)](#) and [D’Antona et al. \(2016\)](#) suggested that either a relatively large variation in He, even by $\Delta Y \sim 0.10$, or a small variation in metallicity, by ~ 0.10 dex, within 1G stars could reproduce the morphology of the 1G sequence on the map. The homogeneity in the light elements among these stars observed in the ChM of many GCs ([Marino et al. 2019](#), see also, [Cabrera-Ziri et al. 2019](#)) suggests that, if He is responsible for the elongated distribution of 1G stars on the ChM, a nucleosynthesis process occurred that changed the He content but left untouched the elements involved in the standard H-burning, such as C and N (see discussion in [Milone et al. 2018](#)). As we are not aware of the existence of such a mechanism, a small variation in metallicity among 1G stars might be a more plausible solution ([Marino et al. 2019](#)).

An internal variation in helium among 1G stars seems not be supported by recent analysis of horizontal branch stars in M3 ([Tailo et al. 2019](#)). The fact that the first population on the maps is inconsistent with one single group poses serious challenges to our current understanding of multiple stellar populations. The possibility that 1G stars have an internal spread in the overall metallicity would suggest that SNe played a role in the star-formation history of most GCs and that these stellar systems have been able to retain a small amount of material polluted by the SNe. In such a case, either a fraction of the 1G could be formed from material polluted from a previous generation of stars thus indicating that only a fraction of 1G stars constitute the real primordial stellar generation, making the mass-budget problem even more dramatic; or the interstellar medium from which 1G stars formed was inhomogeneous. In contrast, the possibility that the 1G has an internal spread in helium (but constant C, N, O, and Fe) could be indicative of some still-unknown mechanism that occurs in GC stars only, possibly an accretion phenomenon in the pre-main sequence phase.

In this work we provide a high-resolution spectroscopic abundance analysis of 1G stars, as selected from the ChM, in the GC NGC 3201. This cluster, displaying a well-elongated 1G population on the ChM, is an

¹ See Section 2. A detailed definition of $\Delta_{F275W,F814W}$ and $\Delta_{C\ F275W,F336W,F438W}$ can be found in [Milone et al. \(2015, 2017\)](#).

ideal target for an in-depth investigation of the chemical abundance pattern within 1G stars. Although its ChM does not show evidence of additional stellar populations on the red side, as in Type II GCs, [Dias et al. \(2018\)](#) found multiple (anti)correlations in light elements, similar to those observed in some of the Type II clusters, such as M 22 ([Marino et al. 2009](#)). The layout of this paper is as follows: Section 2 presents the photometric and spectroscopic data; Section 3 describes how we derive atmospheric parameters; the chemical abundance analysis is presented in Section 4; Sections 5, 6, and 7 discuss our results, that are summarised in Section 8.

2. DATA

2.1. The photometric dataset: the chromosome map of NGC 3201

The photometric data used in this study come from the *HST* UV Legacy Survey which investigated multiple stellar populations in GCs (GO-13297, [Piotto et al. 2015](#)). Details on the images analyzed and on the data reduction can be found in [Piotto et al. \(2015\)](#) and [Milone et al. \(2017\)](#). Photometry has been corrected for differential reddening effects as in [Milone et al. \(2012b\)](#).

[Milone et al. \(2017\)](#) analyzed the chromosome maps of 57 GCs. While we refer to [Milone et al. \(2015, 2017\)](#) for a detailed discussion on how to construct these photometric diagrams, for the convenience of the reader, we restate here their equations (1) and (2). Hence, $\Delta_{F275W,F814W}$ and $\Delta_{C\ F275W,F336W,F438W}$ are defined as:

$$\Delta_{F275W,F814W} = W_{F275W,F814W} \frac{X - X_{\text{fiducial R}}}{X_{\text{fiducial R}} - X_{\text{fiducial B}}} \quad (1)$$

$$\Delta_{C\ F275W,F336W,F438W} = W_{C\ F275W,F336W,F438W} \frac{Y_{\text{fiducial R}} - Y}{Y_{\text{fiducial R}} - Y_{\text{fiducial B}}} \quad (2)$$

where $X = (m_{F275W} - m_{F814W})$, $Y = C_{F275W,F336W,F438W}$, with the latter being an index sensitive to N abundances, and 'fiducial R' and 'fiducial B' correspond to the red and the blue fiducial lines involved in the construction of the ChM.

The analysis of the NGC 3201 ChM revealed that this cluster covers a large range along the $\Delta_{F275W,F814W}$ axis, which is not consistent with the assumed chemical homogeneity of the 1G (see their Fig. 4 and our Fig. 1).

2.2. The spectroscopic dataset

Our spectroscopic data have been acquired using the FLAMES Ultraviolet and Visual Echelle Spectrograph (FLAMES-UVES, [Pasquini et al. 2000](#)) on the European Southern Observatorys (ESO) Very Large Telescope (VLT), through the program 0101.D-0113(A). The observations were taken in the standard RED580 setup,

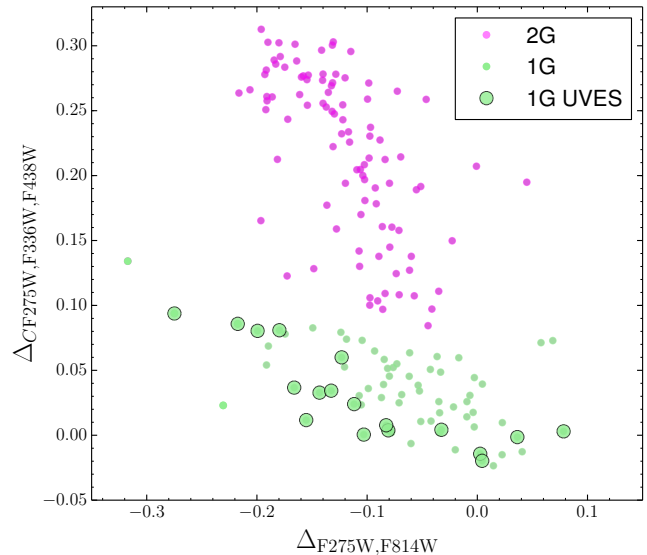


Figure 1. The chromosome map of NGC 3201 from [Milone et al. \(2017\)](#). Stars in green and magenta represent 1G and 2G stars, respectively, as defined by Milone and co-workers. Our 18 spectroscopic UVES targets are all located in the 1G region, and are highlighted as larger filled circles.

which has a wavelength coverage of 4726-6835 Å and a resolution $R \sim 47,000$ ([Dekker et al. 2000](#)).

Spectra are based on $5 \times 2775s$ exposures for the seven brightest stars in our sample with V mag between ~ 13.5 and ~ 14.5 , and $23 \times 2775s$ for the three faintest ones with $V > 15$. Stars with intermediate mag have $11 \times 2775s$ or $12 \times 2775s$ exposures on them (see Table 1). Data were reduced using the FLAMES-UVES pipelines within the EsoReflex interface² ([Ballester et al. 2000](#)), including bias subtraction, flat-field correction, wavelength calibration, sky subtraction and spectral rectification. Once individual spectra were reduced, the telluric subtraction has been performed by using the ESO MOLECFIT tool ([Smette et al. 2014; Kausch et al. 2014](#)).

Radial velocities (RVs) were derived using the IRAF@FXCOR task, which cross-correlates the object spectrum with a template. For the template we used a synthetic spectrum obtained through MOOG³ (version June 2014, [Snedden 1973](#)), computed with a model stellar atmosphere interpolated from the [Castelli & Kurucz \(2004\)](#) grid, adopting parameters (effective temperature/surface gravity/microturbulence/metallicity) = $(T_{\text{eff}}/\log g/\xi_t/[A/H]) = (4900 \text{ K}/2.0/2.0 \text{ km s}^{-1}/-1.50)$. Each spectrum was corrected to the restframe system, and observed RVs were then corrected to the heliocentric system. The mean RVs, together

² <https://www.eso.org/sci/software/esoreflex/>

³ <http://www.as.utexas.edu/chris/moog.html>

with the associated rms obtained from the average of individual exposures, are listed in Table 1. Since NGC 3201 has a distinct high RV, we can safely assume all the observed 18 stars are cluster members. The final mean heliocentric RV for our NGC 3201 giants is $\langle \text{RV} \rangle = +495.3 \pm 0.9 \text{ km s}^{-1}$ ($\sigma = 3.7 \text{ km s}^{-1}$), which lies within 1σ of the literature value of $\langle \text{RV} \rangle = +494.0 \pm 0.2 \text{ km s}^{-1}$ listed in the Harris catalogue (Harris 2010).

Finally, the individual exposures for each star have been co-added. The typical signal-to-noise ratio for the combined spectra around the [O I] $\lambda 6300 \text{ \AA}$ line ranges from $S/N \sim 120$ to ~ 170 , depending on the brightness of the star and the number of exposures. The list of observed giants in NGC 3201 is reported in Table 1, together with coordinates, photometric information, radial velocities (RVs), and number of exposures.

3. MODEL ATMOSPHERES

The relatively high resolution and the large spectral coverage of our spectra allowed a fully-spectroscopic estimate of the stellar parameters, T_{eff} , $\log g$, $[A/H]$ and ξ_t . Hence, we determined T_{eff} by imposing the excitation potential (E.P.) equilibrium of the Fe I lines and gravity with the ionization equilibrium between Fe I and Fe II lines. Note that for $\log g$ we impose Fe II abundances that are 0.05-0.07 dex higher than the Fe I ones to adjust for non-local thermodynamic equilibrium (non-LTE) effects (Bergemann et al. 2012; Lind et al. 2012). This difference is justified by the non-LTE corrections to Fe I and Fe II derived for stars #69 and #82 (see Section 7). The final Fe abundances are based on Fe I, for which we have more available lines. For this analysis, ξ_t was set to minimize any dependence of Fe I abundances as a function of equivalent width (EW).

As an independent test of our results, we also derived atmospheric parameters from our *HST* photometry (see Section 2.1). For that purpose, the m_{F438W} and m_{F606W} magnitudes have been converted to B and V (Anderson et al. 2008), which we then used to estimate temperatures from the Alonso et al. (1999) color-temperature calibrations, assuming a mean $E(B - V) = 0.24$, and a mean $[A/H] = -1.54$ dex. Surface gravities were then obtained from the apparent V magnitudes, the photometric T_{eff} , bolometric corrections from Alonso et al. (1999), apparent distance modulus of $(m - M)_V = 14.20$ (Harris 2010), and a stellar mass of $0.70 M_{\odot}$. Once T_{eff} and $\log g$ have been fixed from photometry, we derived ξ_t from the Fe I lines as explained above.

The top panels of Figure 2 show the spectroscopic targets on the *HST* CMD (left) and the $V - (B - V)$ one (right), with stars colored blue and red for

$\Delta_{F275W, F814W} \leq -0.105$ and $\Delta_{F275W, F814W} > -0.105$, respectively. The small colour offset suggests that, at a given luminosity, the blue stars might be slightly hotter. Our adopted $\log g - T_{\text{eff}}$ values, that are completely independent from photometry do not show significant evidence for such a trend (left-middle panel), which might be slightly more visible in the photometric $\log g - T_{\text{eff}}$ plane (right-middle panel). Our internal errors in T_{eff} is probably larger than the expected difference between bluer and redder stars.

To have an estimate of the internal errors associated with our adopted parameters we compare the $T_{\text{eff}}/\log g$ from Fe lines with those derived from the $(B - V)$ colours in the bottom panels of Figure 2. We obtain: $\Delta T_{\text{eff}} = T_{\text{eff, Fe lines}} - T_{\text{eff, (B-V)}} = +96 \text{ K}$ (rms=47 K), and $\Delta \log g = \log g_{\text{Fe lines}} - \log g_{(B-V)} = +0.01$ (rms=0.10). This comparison suggests that the spectroscopic T_{eff} scale is systematically higher by $\sim 100 \text{ K}$ (see bottom-left panel), but the internal errors are smaller, comparable with the rms of the average differences, i.e. about 50 K. The spectroscopic and photometric gravity scales agree, with a rms of 0.10 dex (bottom-right panel). In the following we adopt typical internal uncertainties of 50 K for T_{eff} , 0.20 dex for $\log g$, 0.20 km s^{-1} for ξ_t , and 0.10 dex for metallicity, but emphasise that systematic errors might be significantly larger. The adopted atmospheric parameters obtained from spectroscopy are listed in Table 2, where we also list the inferred Fe abundances and the $T_{\text{eff}}/\log g$ values from photometry.

4. CHEMICAL ABUNDANCES ANALYSIS

In this work we infer chemical abundances for 24 elements, namely Li, O, Na, Mg, Al, Si, Ca, Sc (II), Ti (I and II), V, Cr (I and II), Mn, Fe (I and II), Co, Ni, Cu, Zn, Y (II), Zr (II), Ba (II), La (II), Pr (II), Nd (II), Eu (II). Chemical abundances were derived from a LTE analysis by using the MOOG code (Snedden 1973), and the alpha-enhanced Kurucz model atmospheres of Castelli & Kurucz (2004), whose parameters have been obtained as described in Section 3. A list of our analyzed spectral lines, with their associated equivalent widths (EWs), excitation potentials (EPs) and total oscillator strengths ($\log gf$), is provided in Tab. 8.

The chemical abundances for all the elements, with the exception of those discussed below, have been inferred from an EW-based analysis. We now comment on some of the transitions that we used.

Lithium: Lithium abundances could be derived for 11 out of the 18 observed giants by spectral synthesis of the Li I $\lambda 6707 \text{ \AA}$ blend using linelist from Lind et al. (2009a) and Hobbs, & Thorburn (1997). The abundances have

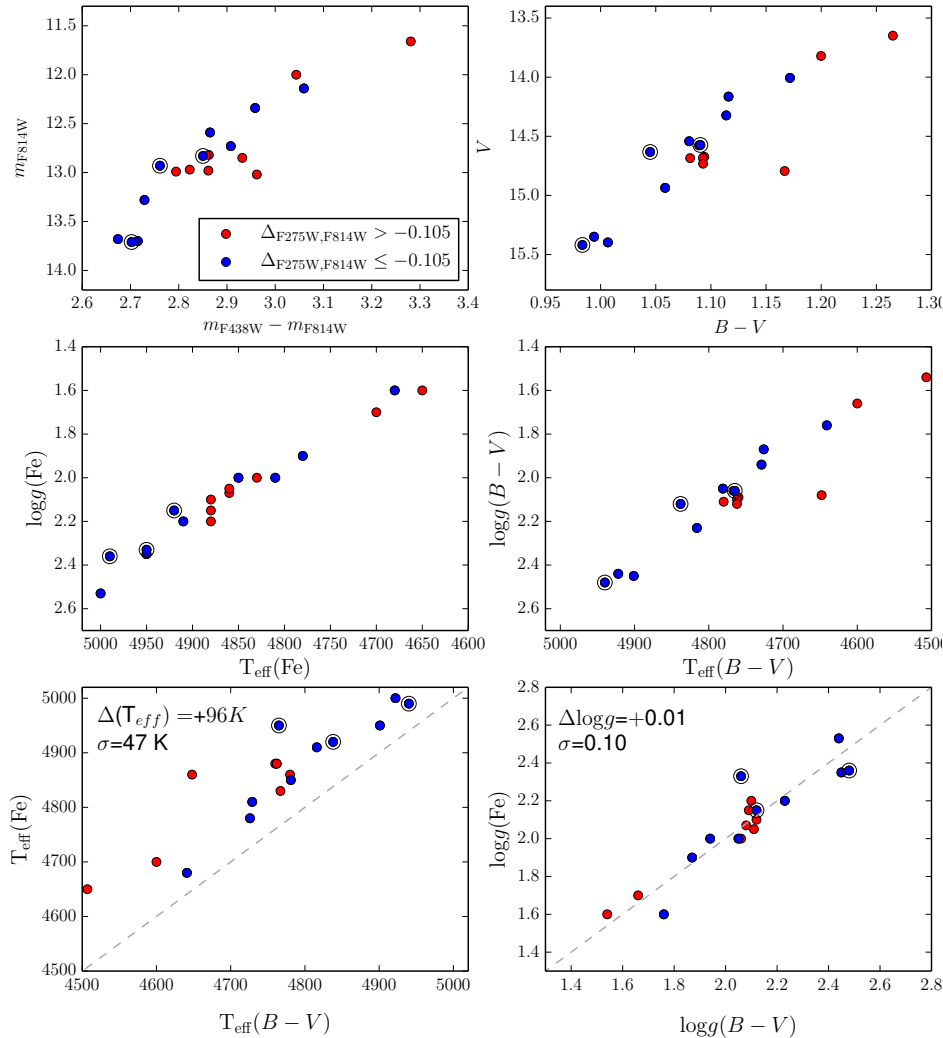


Figure 2. *Upper panels:* CMDs from *HST* photometry. On the left we display the m_{F814W} vs. $m_{F438W} - m_{F814W}$ CMD, while on the right-hand panel we show the V - $(B - V)$ CMD with mag obtained from transforming the original *HST* mags (details in Section 2.1). *Middle panels:* $\log g$ vs. T_{eff} as obtained from the Fe lines (left), and from the $(B - V)$ color (right). *Bottom panels:* Comparison between spectroscopically and photometrically-derived parameters, T_{eff} and $\log g$. The dashed line represents perfect agreement. We label the mean difference and the rms among the two sets of parameters. In all the panels blue and red dots represent stars with $\Delta_{F275W,F814W} \leq -0.105$ and $\Delta_{F275W,F814W} > -0.105$, respectively.

been then corrected for non-LTE effects following Lind et al. (2009a). The location on the color-magnitude diagram clearly suggests that the stars without detectable Li are brighter than the bump, suffering indeed from strong Li depletions that occur at this luminosity. According to the RGB bump V mag provided by Nataf et al. (2013) 2 of the 11 stars with Li measurements are fainter than the bump, hence their abundances might not be directly compared with those in *pre-bump* giants (see Table 3).

Proton-capture elements: In this group of elements, we have derived abundances for O, Na, Al and Mg. Oxygen abundances were inferred from the synthesis of the forbidden [O I] line at 6300 Å. Telluric O₂ and H₂O spectral

absorptions often affect the O line at 6300 Å. Although telluric features have been removed as detailed in Section 2.2 even with such a subtraction procedure, we caution that residual telluric feature contamination might be of concern for the analysis of the 6300.3 [O I] line. We determined Na abundances from the EWs of the Na I doublets at ~ 5680 Å and ~ 6150 Å, aluminum from the synthesis of the doublet at ~ 6667 Å, and magnesium abundances from the EWs of the transitions at ~ 5528 , 5711 Å. Given the weakness of the Al doublet, for four stars we could provide only an upper limit. Sodium abundances have been corrected for deviations from LTE (Lind et al. 2011).

Manganese: For Mn, we have synthesised the spectral lines at around 5395, 5420, 5433, 6014, 6022 Å, by assuming $f(^{55}\text{Mn})=1.00$. When available, the hyperfine splitting data have been taken from Lawler et al. (2001a,b), otherwise from the Kurucz (2009) compendium⁴.

Copper: Abundances for Cu were inferred from synthesis of the Cu I line at around 5105 Å. Both hyperfine and isotopic splitting were included in the analysis, using the well-studied spectral line component structure from Kurucz (2009). Solar-system isotopic fractions were assumed in the computations: $f(^{63}\text{Cu})=0.69$ and $f(^{65}\text{Cu})=0.31$.

Neutron-capture elements: In the group of the neutron-capture (*n*-capture) elements we derived abundances for Y, Zr, Ba, La, Pr, Nd, and Eu. For most of these elements we performed a spectral synthesis analysis, as hyperfine and/or isotopic splitting and/or blending features needed to be taken into account.

Specifically, spectral synthesis was employed for Zr ($\lambda 5112$ Å), Ba ($\lambda 5854$, 6142, and 6497 Å), La ($\lambda 4921$, 5115, 5291, 5304, 6262, 6390, 6774 Å), Pr ($\lambda 5323$ Å), and Eu ($\lambda 6645$ Å). Our Ba abundances were computed assuming the McWilliam (1998) *r*-process isotopic composition and hyperfine splitting. In all the other cases we have assumed the Solar-system isotopic fractions.

The inferred chemical abundances are listed in Tables 3–5. Internal uncertainties to these abundances due to the adopted model atmospheres were estimated by varying the stellar parameters, one at a time, by the amounts estimated in Section 3, namely $T_{\text{eff}}/\log g/[A/H]/\xi_t = \pm 50 \text{ K}/\pm 0.20 \text{ cgs}/\pm 0.10 \text{ dex}/\pm 0.20 \text{ km s}^{-1}$. In addition to the contribution introduced by internal errors in atmospheric parameters, we estimated the contribution due to the limits of our spectra, e.g. due to the finite S/N that affects the measurements of EWs and the spectral synthesis.

To estimate the contribution to the internal uncertainties given by the quality of the spectra ($\sigma_{\text{EWs}/\text{fit}}$), we have compared the EWs obtained from individual exposures of the same stars. For this comparison we used a relatively bright star (#58), with five exposures, and a fainter star (#158), with eleven exposures. We get a typical error in EWs of $\sim 1.5 \text{ m}\text{\AA}$, obtained as the average rms of the EWs measurements for each line divided by the $\sqrt{(N-1)}$, where N is the number of the exposures. For each element, the errors in chemical abundances due to the EWs have been calculated by varying the EWs of spectral lines by the corresponding

uncertainty. For the species inferred from spectral synthesis we have varied the continuum at the $\pm 1 \sigma$ level, and re-derived the chemical abundances from each line.

For the spectral lines analysed with synthesis we follow the approach by Norris et al. (2010) and Yong et al. (2013b). For each element, we replace the rms (σ) in Tables 3–5 by the maximum σ value. Then, we derive $\max(\sigma)/\sqrt{N_{\text{lines}}}$. For those elements whose abundance is inferred from just one line we use the typical uncertainty introduced by the continuum scatter as our $\sigma_{\text{EWs}/\text{fit}}$. Typical values obtained for each element are listed in column 8 of Table 6. The total error is obtained by quadratically adding this random error with the uncertainties introduced by atmospheric parameters.

Since the EWs/continuum placement errors are random, the uncertainty associated to those elements with a larger number of lines is lower. Hence, the corresponding uncertainty associated with Fe I is negligible, while for those species inferred from one or two weak spectral lines, the error due to the limited S/N is dominant (e.g. O, Al, Zr, Pr, and Eu).

5. THE CHEMICAL COMPOSITION ALONG THE FIRST POPULATION OF NGC 3201

In this section we explore the chemical composition in different elemental species along the elongated 1G observed on the ChM of NGC 3201. We emphasize that, although the range covered by our targets in the $\Delta_{F275W, F814W}$ axis is large, they span only a small range in $\Delta_{C F275W, F336W, F438W}$, and their position on the map is consistent with the 1G, as defined by Milone et al. (2015, 2017). As the $\Delta_{C F275W, F336W, F438W}$ axis of the ChM is shaped by light elements variations, primarily driven by the N enhancements in the 2G stars, we expect homogeneous abundances in this chemical species (Marino et al. 2019).

Overall, for our 18 giants we obtain a mean iron abundance of $[\text{Fe}/\text{H}] = -1.50 \pm 0.02 \text{ dex}$ (rms=0.07 dex), consistent with the value of $[\text{Fe}/\text{H}] = -1.59$ listed in Harris (2010). Figure 3 shows a summary of the other chemical abundance ratios obtained for our 1G sample of NGC 3201 stars⁵. Besides observing the typical chemical pattern of Population II stars, namely the enhancement in the α -elements (Mg, Si, Ca, Ti), and the typical solar-scaled abundances of Fe-peak elements, more interesting for this study is the investigation of the elements that have a close connection with the shape of the ChM.

⁵ Chemical abundances are expressed in the standard notation, as the logarithmic ratios with respect to solar values, $[X/Y] = \log(\frac{N_X}{N_Y})_{\text{star}} - \log(\frac{N_X}{N_Y})_{\odot}$. For lithium, abundances are reported as $A(\text{Li}) = \log(\frac{N_{\text{Li}}}{N_{\text{H}}})_{\text{star}} + 12$.

⁴ <http://kurucz.harvard.edu/>

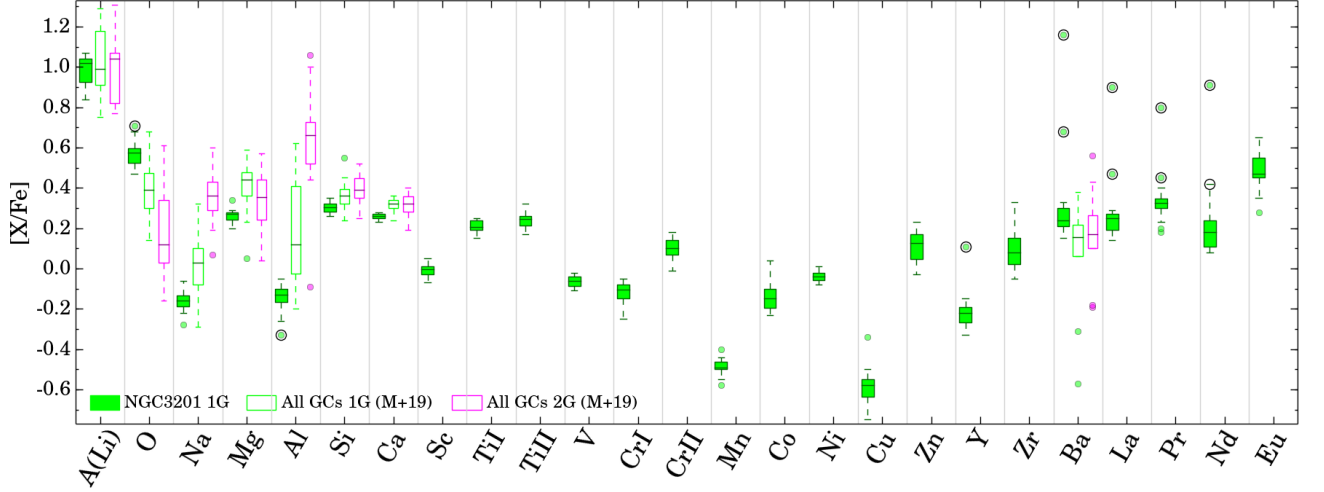


Figure 3. Box-and-whisker plot for the chemical abundances of the 1G stars observed on the ChM of NGC 3201 (green filled boxes). All the abundances are relative to Fe, except for Li, for which we use the usual $A(\text{Li})$ notation. Plotted Li abundances are corrected for non-LTE effects. As a comparison, we also show the box-and-whisker plots for all the GCs for which abundances on ChMs are available. These boxes have been constructed with the average abundances for 1G and 2G listed in Table 2 from Marino et al. (2019). 1G and 2G average abundances for all GCs are represented in green and magenta empty boxes. Each box represents the interquartile range (IQR) of the distribution, with the median abundance marked by an horizontal line. The whiskers include observations that fall below the first quartile minus $1.5 \times \text{IQR}$ or above the third quartile plus $1.5 \times \text{IQR}$. As the majority of the used literature studies for Na report LTE abundances, the plotted $[\text{Na}/\text{Fe}]$ values for NGC 3201 are not corrected for non-LTE. Small filled circles represent outliers in the data. Outliers highlighted with black open circles are the candidate binaries (see Section 6 for details).

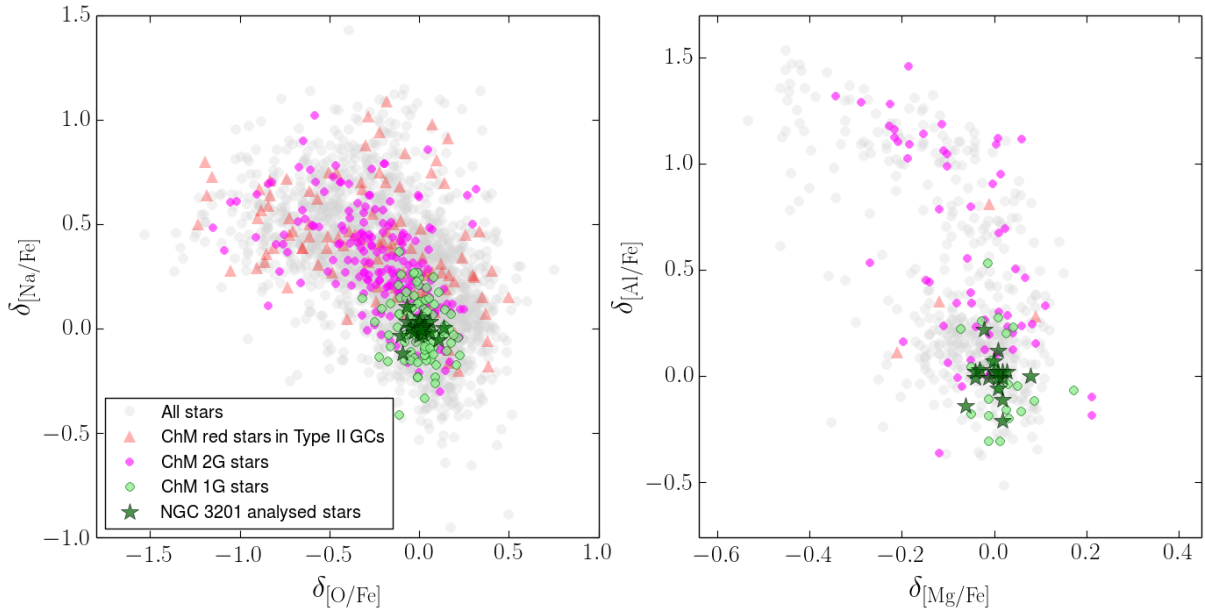


Figure 4. Abundance ratios of $\delta_{[\text{Na}/\text{Fe}]}$ as a function of $\delta_{[\text{O}/\text{Fe}]}$ (left) and $\delta_{[\text{Al}/\text{Fe}]}$ as a function of $\delta_{[\text{Mg}/\text{Fe}]}$ (right) for all the GCs analysed in Marino et al. (2019). The plotted abundances are relative to the average abundances of 1G stars, as defined by Marino and collaborators. Different symbols and colors represent stars belonging to different populations on the ChM. The abundances of the analysed stars in NGC 3201 perfectly overlap with the ChM 1G abundances on these $\delta_{[\text{O}/\text{Fe}]}-\delta_{[\text{Na}/\text{Fe}]}$ and $\delta_{[\text{Mg}/\text{Fe}]}-\delta_{[\text{Al}/\text{Fe}]}$ planes.

In Figure 3 we show a comparison of our results for NGC 3201 with all the GCs for which abundances on ChMs are available from Marino et al. (2019, Table 2). Overall, our inferred abundances agree with those observed in the 1G as selected on the ChM of GCs, as expected from the location of our targets along the ChM of NGC 3201. In the next sections we will discuss all the interesting abundance patterns we observe in NGC 3201, focusing on those elements that play an important role in the multiple stellar populations phenomenon and in shaping the ChM. In particular, Marino et al. (2019) have analysed the chemical abundance pattern along the ChM, for the species most involved in the multiple stellar population phenomenon, namely the p -capture (e.g. O, Na, Mg, Al) and n -capture elements (Ba).

First, we note that the abundances relative to Fe ($[X/Fe]$) are generally consistent with uniform chemical abundances in the plotted elements. By comparing the observed rms associated to the mean average abundances, as listed in Tables 3–5, with the estimated errors (Table 6), it appears that, in most cases, our expected uncertainties are higher. This suggests that our estimated errors might be overestimated. An exception to this general trend are the n -capture elements, that will be discussed in the following sections.

As already mentioned, the chemical abundances of the 18 stars of NGC 3201 analysed here are consistent with the chemical composition of 1G stars. Lithium abundances have a range of ~ 0.2 dex, from $A(\text{Li}) \sim 0.75$ up to $A(\text{Li}) \sim 1.0$ dex, which compares to typical values of RGB stars having experienced the full first dredge up but not having reached yet the RGB bump level (Lind et al. 2009b). As highlighted in Table 3, stars with only upper limits plus two other stars with Li measurements are brighter than the RGB bump. In the context of multiple stellar populations, Marino et al. (2019) did not find strong evidence for variations of Li between 1G and 2G stars on the ChM, with the exception of GCs like NGC 2808, and ω Centauri. Hence, the comparison of Li with the average abundances of 1G and 2G stars in all the GCs, as shown in Figure 3 does not likely provide clear information on the population *status* of our sample in NGC 3201.

Oxygen and sodium are the best tracers of the multiple stellar populations phenomenon and their abundances are indicative of different locations along the ChM. Specifically, for our NGC 3201 giants the oxygen abundances relative to iron are super-solar, as typical of Population II stars, while sodium is not enhanced. Both these elements show distributions compatible with the 1G ones. The association of all the analysed stars with the chemical composition of the ChM 1G is clear

by looking at their location on the $\delta_{[\text{O}/\text{Fe}]} - \delta_{[\text{Na}/\text{Fe}]}$ plane plotted in Figure 4, where the δ abundances are the relative abundances to the average chemical content of 1G stars in each cluster (see Marino et al. 2019, for a detailed discussion).

Chemical enhancements in aluminum, coupled with depletions in magnesium are observed in a few clusters, with NGC 2808 being a noticeable example. Both our Mg and Al abundances fall in the range spanned by 1G stars. Magnesium relative to Fe is super-solar, as expected from a typical α -element, and the Al abundances distribute on the lower abundance tail spanned by 1G stars. Both these element distributions do not show any evidence for internal variations within our 18 analysed stars. The location of our stars on the $\delta_{[\text{Mg}/\text{Fe}]} - \delta_{[\text{Al}/\text{Fe}]}$ plane, constructed by using relative abundances similar to those used for $\delta_{[\text{O}/\text{Fe}]}$ and $\delta_{[\text{Na}/\text{Fe}]}$, are plotted in the right panel of Figure 4.

6. BINARITY AND BLUE STRAGGLERS

Noticeably, Figure 3 shows a few stars with significantly higher abundances in the n -capture elements, Y, Ba, La, Pr and Nd. This class of chemical species plays an important role in the Type II GCs, those displaying additional ChMs on the red side of the main map. Stars on the red side of ChMs are typically enhanced in both Fe and the elements mostly produced via s -process reactions, e.g. Ba and La, but not in Eu, a typical r -process element. Nevertheless, NGC 3201 does not belong to the class of Type II GCs, as it does not show any clear additional redder sequence in the ChM as M 22, M 2, or ω Centauri do (Milone et al. 2017).

To investigate the nature of the variations in s -process elements in NGC 3201, we plot in Figure 5 the abundance of $[\text{La}/\text{Fe}]$ taken as representative of the s -process elements, as a function of the r.m.s. in the RVs (left-hand panel) and $\Delta_{\text{F275W}, \text{F814W}}$ (right-hand panel). While most stars show typical $[\text{La}/\text{Fe}]$ abundances around 0.25 dex, two stars have $[\text{La}/\text{Fe}] > 0.40$ dex, with one reaching an abundance as high as $[\text{La}/\text{Fe}] = 0.90$ dex. Interestingly, this extremely-La rich star (#93) also has a high r.m.s. in RV (σ_{RV}), compared to the bulk of the stars lying around 0.2 km s^{-1} . A third outlier in the $[\text{La}/\text{Fe}] - \sigma_{\text{RV}}$ plane displays the highest r.m.s. in RV (0.76 km s^{-1} , star #67), but it is not an outlier for $[\text{La}/\text{Fe}]$.

The high RV r.m.s. in two giants of our sample suggests that these stars are indeed binaries. By applying a maximum likelihood approach, similar to that used in Frank et al. (2015) and Piatti, & Koch (2018), assuming a typical uncertainty of 0.17 km s^{-1} , we found that the intrinsic dispersions in the RVs of the stars #93 and

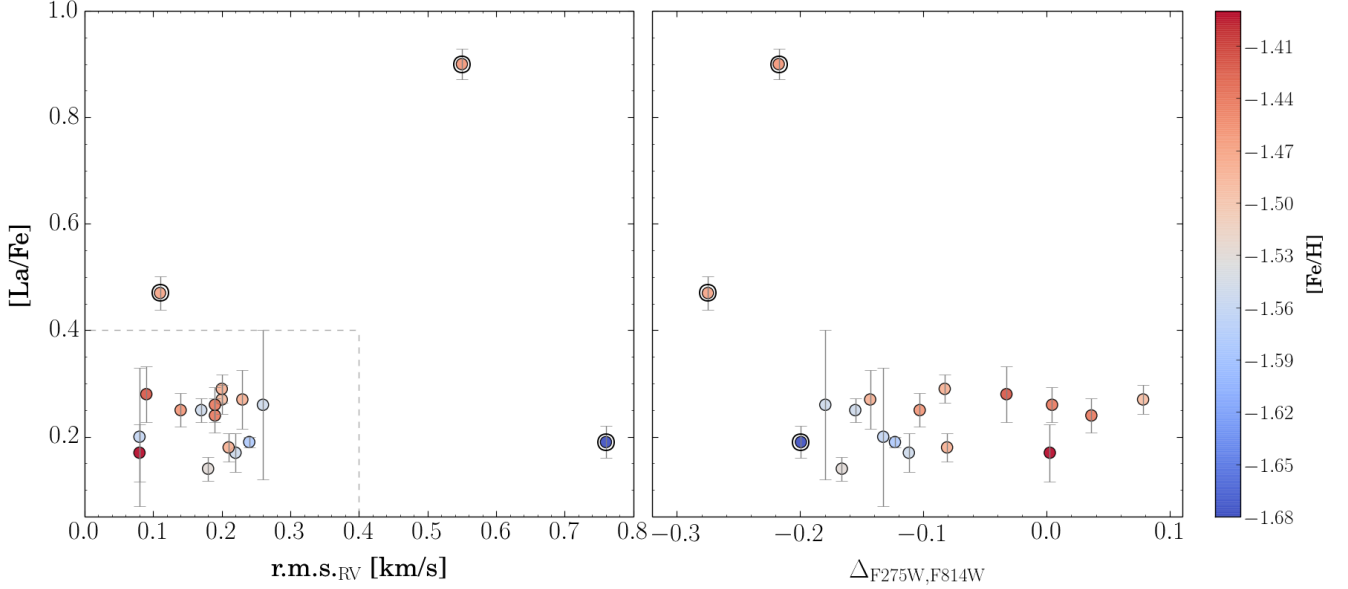


Figure 5. The abundance of the n -capture element La relative to Fe, as a function of the r.m.s. in the RVs (left panel), and $\Delta_{F275W,F814W}$ (right panel). Error bars are the statistical errors associated to the average abundances for each star obtained from different lines. The color of each star is indicative of its inferred Fe, as illustrated in the right-side colour bar. Our three binary candidates, outlined with black circles, lie outside the box delimited by a dashed line in the left panel.

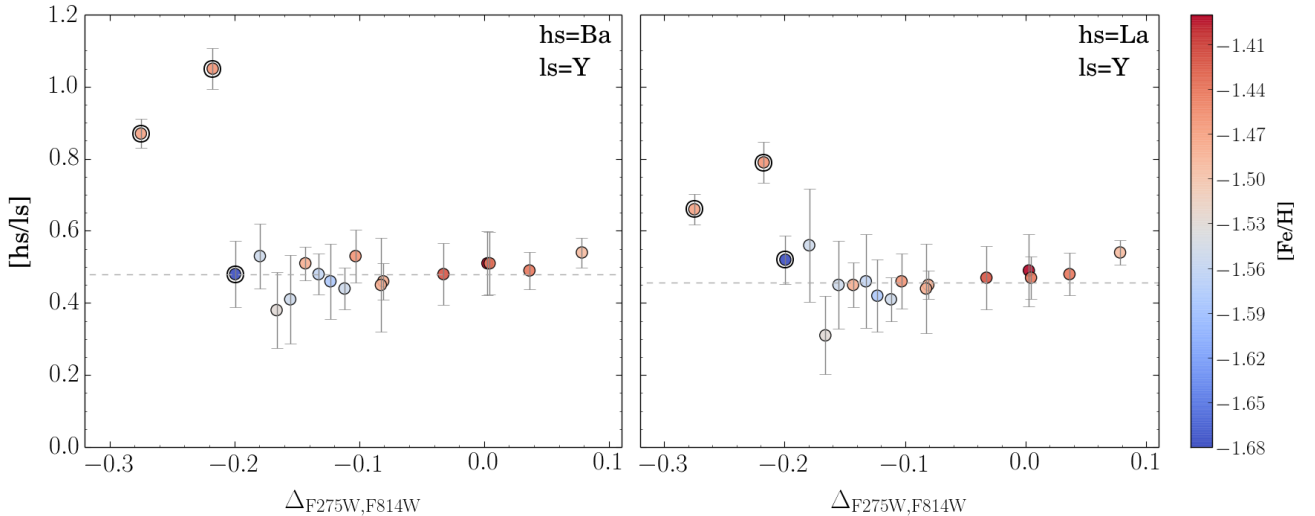


Figure 6. The abundance of [Ba/Y] (left) and [La/Y] (right), representative of the heavy to light s -element abundances ([hs/ls]), as a function of $\Delta_{F275W,F814W}$. The dashed lines in both panels highlight the average abundance ratios for the entire sample, neglecting the three (encircled) binary candidate. As in Figure 5, the color of each star is indicative of its inferred Fe, as illustrated in the right-side colour bar.

#67 are 0.51 and 0.74 km s^{-1} , respectively. Such values, higher than zero, corroborate the idea that these stars are binary candidates. The very high abundance in s -elements of one of these stars is consistent with the binary nature of this object. Similarly, star #149 having relatively high abundances in s -elements can be regarded as a binary candidate, although our data do not show any evidence of high variations in its RV.

As shown in Figure 6, the stars with high n -capture element abundances also show higher [hs/ls] than the mean value of ~ 0.5 dex, suggesting that they are likely the result of direct mass transfer with a low-mass asymptotic giant branch (AGB) companion. The abundance ratios [La/Y] and [Ba/Y], as a function of $\Delta_{F275W,F814W}$, are indicative of the heavy to light n -capture element abundance ratio, which is sensitive to the neutron exposure and neutron density.

Higher $[\text{hs}/\text{ls}]$ is indicative of the operation of the $^{13}\text{C}(\alpha,n)^{16}\text{O}$ neutron source operating in low-mass AGB stars (e.g., Gallino et al. 1998; Fishlock et al. 2014), where values around zero or negative values indicate the $^{22}\text{Ne}(\alpha,n)^{25}\text{Mg}$ reaction. The higher $[\text{hs}/\text{ls}]$ of the two giants enriched in n -capture elements might suggest direct mass transfer between the star we now observe and previous low-mass, low-metallicity AGB stars (e.g. Cristallo et al. 2009; Karakas, & Lattanzio 2014).

In the last column of Table 1 we list the maximum time interval between the observations of each star. We note that some stars, including #149, have been observed over a shorter time interval (3 days rather than months), which might have prevented us from detecting RV measurable variations in longer-period binaries.

In the following we will consider these three stars as binary candidates. Their locations in Figures 2, 3, 5 (as well as in some of the following figures) are highlighted with black open circles. We warn the reader that we cannot rule out the presence of other binaries with no obvious s -elements or/and RV differences among our remaining 15 giants.

Another stellar population that may affect the appearance of the chromosome maps, and likely has a connection to binary stars, are the blue stragglers (BSs). In particular, we consider their evolved counterparts which lie close to the red giant branch. In Figure 7 we use blue symbols to represent the candidate BS stars of NGC 3201 in the $m_{\text{F}336\text{W}}-(m_{\text{F}275\text{W}} - m_{\text{F}814\text{W}})$ CMD. The objects marked with blue starred symbols are possible BSs, selected on the CMD, that are evolving towards the RGB phase, which clearly display lower $C_{\text{F}275\text{W},\text{F}336\text{W},\text{F}438\text{W}}$ values in the $m_{\text{F}336\text{W}}-C_{\text{F}275\text{W},\text{F}336\text{W},\text{F}438\text{W}}$ diagram (upper-right panel). Blue stragglers in GCs are tightly linked to the cluster binary populations (Knigge et al. 2009). Blue stragglers in open clusters themselves have a very high binary fraction (Mathieu, & Geller 2009), and preliminary evidence suggests the same is true in NGC 3201 (Giesers et al. 2019). When we include the evolved BSs in the ChM (lower panel of Figure 7), they populate a well-defined sequence on the blue extension of the 1G sequence. We conclude that, based on RVs and s -process element abundance, the three bluest 1G stars that we analyzed spectroscopically are binary systems. Some of them are possibly associated with the BS population of NGC 3201, as evolved BSs would contribute to the bluest extension of the 1G sequence in the ChM.

6.1. Non-interacting binaries simulations

To investigate the effect of non-interacting binaries formed by pairs of 1G stars, we reproduce in Figure 8

five isochrones, I1—I5, with $[\text{Fe}/\text{H}]=-1.50$, $[\alpha/\text{Fe}]=0.4$ and age of 13 Gyr from Milone et al. (2018). These isochrones are constructed with a different combination of He, C, N, and O, to represent the typical chemical pattern of different stellar populations in GCs. We used the isochrone, I1, which corresponds to the 1G, to generate a stellar population of 1G-1G binaries that we represent with black points in Figure 8.

The corresponding ChM plane of RGB stars for the five populations, is shown on the right panel. Together with the five isochrones with different chemical abundances, we plot a population of simulated 1G-1G binaries (black points) constructed by assuming a flat distribution in the mass ratio (q). This test clearly demonstrates that binaries can contribute to the width of the 1G in the ChM. However, we find that the shift towards bluer $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ is dependent on the mass ratio and only binaries with large mass ratio ($q \gtrsim 0.8$) are able to provide a significant color spread.

To have a more direct comparison with the ChM of NGC 3201, in Figure 9 we superimpose to the observed map two simulated binary populations: (i) a population of binaries that account for the whole cluster (100% of binaries), with both the 1G-1G and the 2G-2G pairs (upper panel); (ii) a more realistic simulation, where the binary population fraction is 12.8%, as reported in Milone et al. (2012b) for the main sequence binaries (lower panel). In the latter we assume that all the binaries are 1G-1G pairs. Again we find that the 1G-1G binaries can affect the elongation of the 1G on lower $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ values, but the observed binary fraction on the main sequence are not able to account for the relatively high number of stars observed on the bluer extension of the ChM.

7. CHROMOSOME MAP AND ABUNDANCE PATTERN OF 1G STARS

On the right-hand panel of Figure 5, we show the $[\text{La}/\text{Fe}]$ abundances as a function of the $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ axis of the ChM. The three binary candidates are the stars with the lowest $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ values in our sample. As discussed in the previous section, this suggests that the stars with the lowest $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ on the map are likely associated with binaries and/or BSs, with no need to invoke chemical abundances variations like He enhancements.

However, our simulations suggest that we would need a high number of binaries at high q to account for all the stars in the 1G population. This requirement might suggest that some additional mechanism might be necessary to account for a range of ~ 0.2 mag in $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$

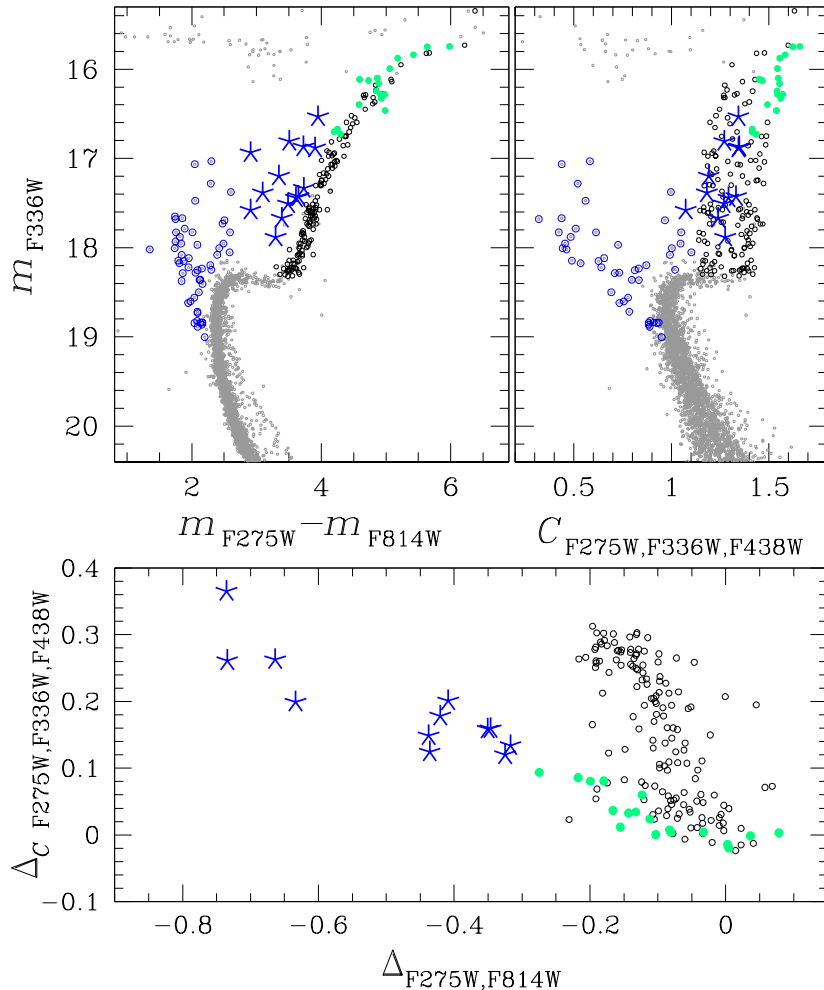


Figure 7. *Upper panels:* $m_{F336W} - (m_{F275W} - m_{F814W})$ CMD (left) and $m_{F336W} - C_{F275W, F336W, F438W}$ diagram (right panel) of NGC 3201. *Lower panel:* ChM of NGC 3201, similar to Figure 1, but extended to show the evolved BSSs. In all the panels the green dots represent our analysed spectroscopic targets, the open blue circles, the blue stragglers, and the star-like symbols are the evolved blue stragglers.

among 1G stars in NGC 3201, and similarly in many other GCs.

In Figure 5 the used color code is indicative of the Fe abundance. We immediately note that the stars with lower $\Delta_{F275W, F814W}$ also have lower iron, with #27 ($\Delta_{F275W, F814W} = +0.0026$) and #33 ($\Delta_{F275W, F814W} = -0.1231$) being the stars with the highest and lowest metallicity, respectively, giving a maximum range of 0.19 dex (neglecting the three binary candidates).

This result suggests that there is a small Fe variation among 1G stars in NGC 3201. By dividing our sample of 15 stars on the basis of the $\Delta_{F275W, F814W}$ value, we find that eight stars with $\Delta_{F275W, F814W} < -0.10$ have $\langle [\text{Fe}/\text{H}] \rangle = -1.53 \pm 0.02$, while the seven stars at higher $\Delta_{F275W, F814W}$ have $\langle [\text{Fe}/\text{H}] \rangle = -1.45 \pm 0.01$, more than 2.5σ difference. This difference is smaller than that inferred for the blue- and red-RGB stars in the Type II GCs (Marino et al. 2019), and hardly detectable. Our

relatively high-S/N and high-resolution UVES spectra for 15 stars allow us to detect the difference still at a $< 3\sigma$ level. We also note that the dispersion associated with our $[\text{Fe}/\text{H}]$ average abundance is comparable with the estimated error.

It is worth, at this point, to discuss the impact of non-LTE corrections on this result, which is based on LTE abundance analysis. We first note that, as amply discussed in Kovalev et al. (2019), non-LTE analysis changes the mean abundance ratios in clusters, but intra-cluster abundance dispersions should not be significantly affected. However, to test this we derived the non-LTE corrections to our Fe I and Fe II spectral lines (Bergemann et al. 2012; Lind et al. 2012), by using the online available database (<http://inspect-stars.com/>). For this test we use stars #69 and #82, that have similar atmospheric parameters (both from spectroscopy and photometry), but different $\Delta_{F275W, F814W}$,

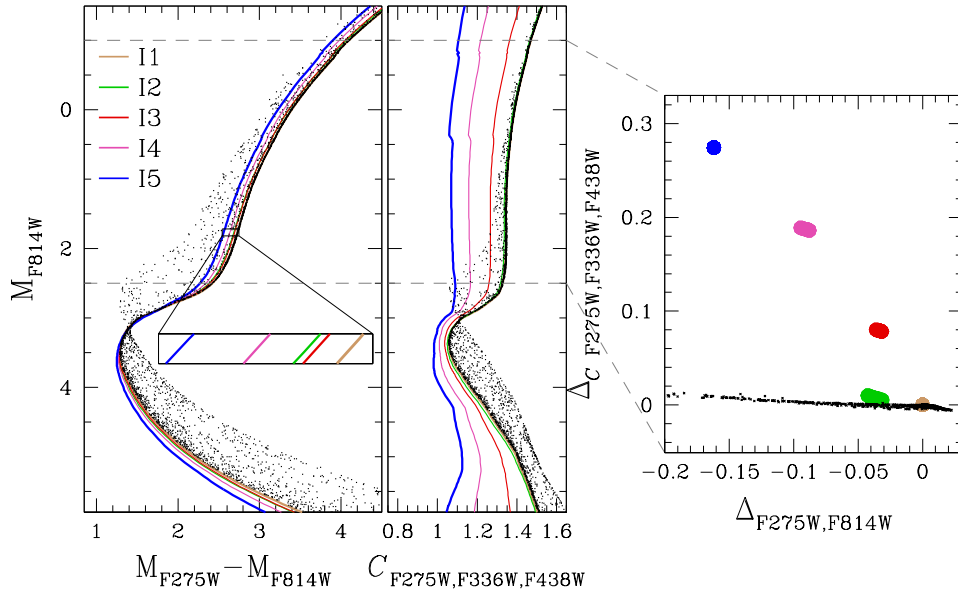


Figure 8. *Left panels:* Simulated CMD, with five isochrones corresponding to different stellar populations with different He, C, N, O: I1 is the isochrone corresponding to the first population; I2 has the same abundance of C, N and O as I1 but different He; in I3, I4, and I5 abundances of He, C, N, O have been varied to represent the second population chemical pattern. *Right panel:* Location of the different stellar simulated five populations at a reference magnitude on the ChM plane. The black points are non interacting binaries formed by pairs of I1-I1 stars. Binaries can contribute to the width of the 1G in the ChM, but only binaries with large mass ratio ($q \gtrsim 0.8$) can provide a significant color spread.

with star #82 having a higher derived $[\text{Fe}/\text{H}]$. We find that the non-LTE correction to apply to Fe I lines is $+0.07$ dex for both stars, and does not change significantly from line to line. Fe II abundances decrease by -0.01 , again in both stars. Using non-LTE Fe abundances for the determination of the atmospheric parameters only marginally changes the spectroscopic temperature, by $+30$ - 40 K in both stars. From this test, it is clear that, although non-LTE abundances are more realistic, in our sample of stars the corrections are systematic and similar in all the stars. The difference in $[\text{Fe}/\text{H}]$ between the two stars is kept unchanged.

Figure 10 shows $[\text{Fe}/\text{H}]$ as a function of $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$. Excluding the three binary candidates we get a positive correlation (slope=0.50), with a Spearman correlation coefficient $r=0.68$. We have derived the significance of the correlation from a Monte Carlo of 1000 realizations of our dataset composed of 15 stars. In each realization we have assumed the observed $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ and a uniform Fe abundance with the associated error estimates of Table 6, and derived the slope. We have calculated the fraction (f) of realizations where the slope is equal or higher than the observed one, and assumed this value as the probability that the slope is due to randomness. For $[\text{Fe}/\text{H}]$ we get a slope higher than the observed value in 2.2% of the simulations (see Table 7).

Figure 11 represents the Spearman correlation for each $\log\epsilon(\text{X})$ abundance as a function of the $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$

and $\Delta_{\text{C F}275\text{W},\text{F}336\text{W},\text{F}438\text{W}}$ value on the ChM of NGC 3201. The obtained correlation values between $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ and the $\log\epsilon$ abundances are listed in Table 7, together with the associated f . Although not all the correlations are significant, given the relatively small sample size, clearly all the abundances appear to be positively correlated with $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$. We regard the negative correlation with $\Delta_{\text{C F}275\text{W},\text{F}336\text{W},\text{F}438\text{W}}$ as simply due to the fact that our stars with low $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ have slightly higher $\Delta_{\text{C F}275\text{W},\text{F}336\text{W},\text{F}438\text{W}}$ (see Figure 1). However, the presence of a general correlation between each absolute abundance and the $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ corroborates the presence of a small spread in the overall metallicity among 1G stars in NGC 3201.

A quick look at the spectra further supports this finding. As an example, in Figure 12 we show the portion of two spectra, including some Fe and Ca analysed spectral features, for two stars with different $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ and similar atmospheric parameters. We note that the two chosen stars also have very similar atmospheric parameters from photometry. Overall, the spectral features of the star at higher $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ look consistent with higher metals.

As first suggested by Milone et al. (2015, 2017) a variation in the He content intrinsic to the 1G stars is qualitatively able to account for the $\Delta_{\text{F}275\text{W},\text{F}814\text{W}}$ spread. However, such a variation in He should be high, at a

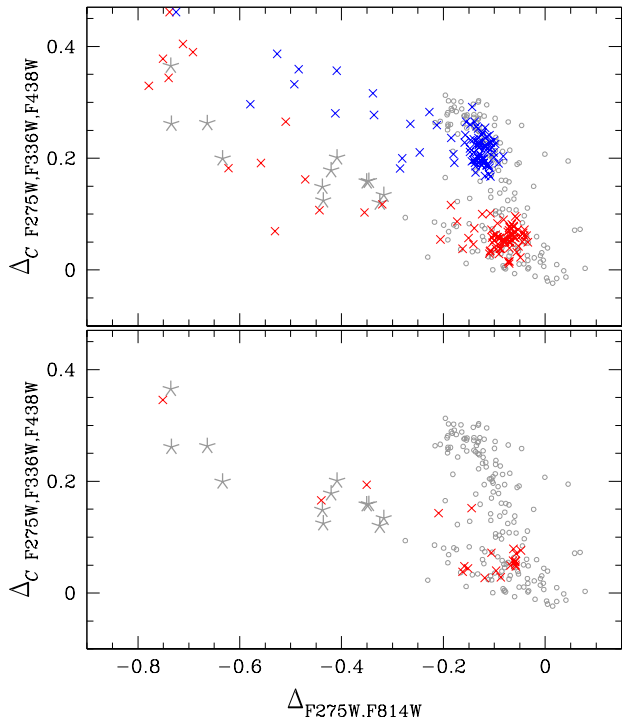


Figure 9. Observed ChM of NGC 3201 (grey symbols). The stars with extremely low $\Delta_{F275W,F814W}$ are plotted as grey star-like symbols. Superimposed to the observed ChM are simulations of binary stars: in the upper panel we have simulated a cluster with the 100% of binaries, both the 1G-1G (red) and the 2G-2G (blue) binaries; in the lower panel, we have simulated the 12.8% of binaries, which is the observed fraction on the main sequence (Milone et al. 2012b), by assuming all of them being 1G-1G binaries.

level of $\Delta Y \sim 0.08-0.10$, which is difficult to achieve without any corresponding enhancement in other chemical species, such as N and Na, and depletion in O (Marino et al. 2019). As discussed in Section 5 our analysed stars in NGC 3201 have similar $[O/Fe]$ and $[Na/Fe]$ abundances.

Assuming the working hypothesis that a He variation exists among our stars, we would expect a difference in the structure of a model atmosphere between a He-normal and a He-rich star (Stromgren et al. 1982). Such a difference would translate in a different surface gravity, which is expected to be small (<0.10 dex). Looking at the impact of a variation in $\log g$ by 0.20 dex on the chemical abundances, it is unlikely that such variation in $\log g$ introduced by a possible He variations are responsible for the abundance variations we find.

As widely discussed in Yong et al. (2013a), a second effect of possible He variations is that for a fixed mass fraction of metals (Z), a change in the helium mass fraction (Y) will directly affect the hydrogen mass fraction (X) such that the metal-to-hydrogen ratio, Z/X will change with helium mass fraction since $X + Y + Z = 1$. Hence,

if stars in a globular cluster have a constant Z , an He-rich star will appear to be slightly more metal-rich than an He-normal star. In this context, by using spectra of excellent quality Yong et al. (2013a) found variations at a level of a few hundredth of dex in the absolute chemical abundances of giants in NGC 6752, positively correlated with Na (a tracer of He)⁶.

For a fixed Z , in the case of a variation in Y of 0.10 (the one predicted to account for the $\Delta_{F275W,F814W}$ extension in NGC 3201), $[Z/H]$ would change by ~ 0.06 dex, which is slightly lower than the range we find in the cluster. However, as a *pure* He enhancement shifts stars towards lower $\Delta_{F275W,F814W}$ values (Milone et al. 2015; Marino et al. 2019), this second explanation can be ruled out for our 1G stars in NGC 3201, as the stars that should be enhanced in He (at lower $\Delta_{F275W,F814W}$), are metal-poorer. Clearly, they are located on the blue side of the RGB, and of the ChM, reinforcing the idea that, at least the bluest side of 1G in the ChM, is populated by binaries and/or evolved blue stragglers.

On the other hand, if we assume that all the metal-poorer and bluest stars are indeed binaries, we expect that the sum of the spectra of the two binary components is reflected in some variation in the line-flux/continuum ratio. Depending on the brightness of the two stars, these effects might not be negligible and can affect the derived atmospheric parameters. To qualitatively investigate this issue, we have simulated some spectra with the SYNTHE routine in the ATLAS code (Kurucz 2009). For this purpose, we considered a bright giant star with $(T_{\text{eff}}/\log g/[A/H]/\xi_t) = (4650 \text{ K}/1.60/-1.50/2.00 \text{ km s}^{-1})$, and two less luminous stars, namely a fainter giant $(5000 \text{ K}/2.53/-1.50/2.00 \text{ km s}^{-1})$, and a sub-giant $(6000 \text{ K}/4.00/-1.50/2.00 \text{ km s}^{-1})$. By summing the flux of the bright giant and the sub-giant, the combined spectrum is almost identical to the giant spectrum. However, the sum of the two giants results in an emerging spectrum which is consistent with either lower overall metallicity (by 0.06 dex) or higher T_{eff} (by 50 K). Binaries where the two components are both giants, but with different luminosities can potentially explain the lower metals in the bluest 1G stars, however as far as we know it is unlikely to have as many giant pairs as the bluer 1G stars (Milone et al. 2012b).

⁶ For NGC 6752, high-precision photometry infers a ΔY of the order of a few hundredth (Milone et al. 2018), not enough to entirely account for the observed abundance correlations derived in Yong et al. (2013a). Yong and co-workers suggest that a combination of He variations and inhomogeneous chemical evolution in the protocluster environment could account for the chemical abundance variations.

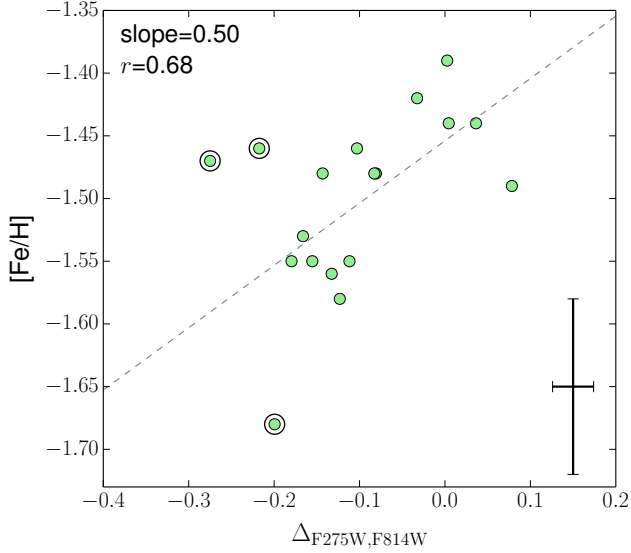


Figure 10. Chemical abundances $[\text{Fe}/\text{H}]$ as a function of $\Delta_{\text{F275W},\text{F814W}}$ values on the ChM of NGC 3201. The dashed line is the least squares linear fit to the data (neglecting the three binary candidates). The slope and the Spearman correlation coefficient (r) are highlighted in the top-left corner.

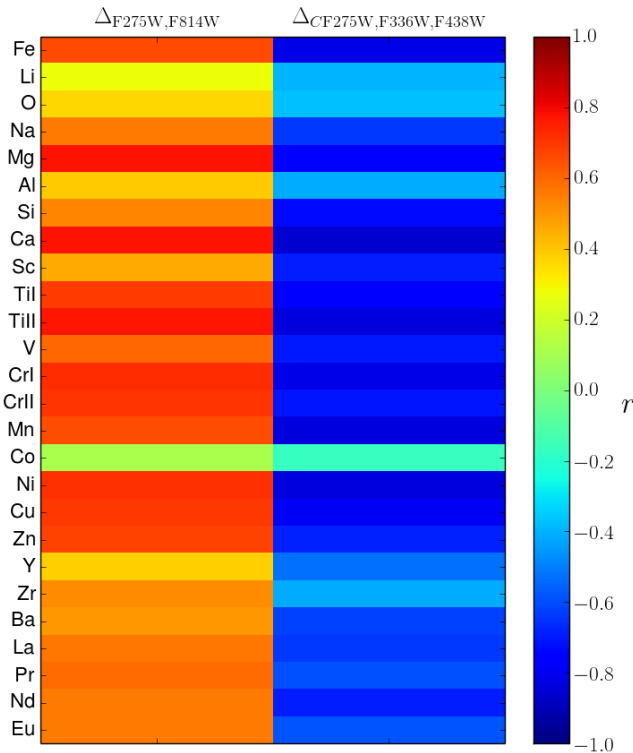


Figure 11. Chemical abundances $\log\epsilon(X)$ as a function of $\Delta_{\text{F275W},\text{F814W}}$ and $\Delta_{\text{CF275W},\text{F336W},\text{F438W}}$ values on the ChM of NGC 3201. For Li and Na we use the non-LTE abundances. The color code is indicative of the Spearman correlation coefficient (r). The significance of these correlations is listed in Table 7.

8. DISCUSSION AND CONCLUSIONS

We have presented a high-resolution chemical abundance analysis of eighteen stars of the GC NGC 3201. The eighteen stars have been selected all to belong to the 1G stellar population, as defined on the ChM diagnostic tool in Milone et al. (2017). Although located on the 1G ChM sequence, our targets span a large range in the $\Delta_{\text{F275W},\text{F814W}}$ axis.

Overall, we have found that the abundances relative to Fe of all the inferred species are consistent with a uniform chemical composition relative to Fe. Specifically, the stars have similar O and Na, consistent with the 1G chemical abundances. This finding confirms previous results showing that the elongation in $\Delta_{\text{F275W},\text{F814W}}$ of the 1G in GCs is not related to the common light-elements (anti)-correlations (Marino et al. 2019; Cabrera-Ziri et al. 2019). The constancy in light elements makes it hard to believe that the stars with lower $\Delta_{\text{F275W},\text{F814W}}$ have higher He abundances, as tentatively suggested by Milone et al. (2015).

The three stars with the lowest $\Delta_{\text{F275W},\text{F814W}}$ values in the sample are binary candidates (see Figure 5). Two of these stars have higher RV r.m.s., over an observation time of a few months, than the bulk of our stars. One star (#93) also shows extremely high s -process element abundances. A high s -element abundance is associated to long-period, single-lined spectroscopic binary systems (McClure 1989). The unseen companion was a low-mass AGB star that transferred processed material to the surviving visible star. A third star shows higher s -element abundances, although not as extreme as #93, but its r.m.s. in RVs is similar to the bulk of stars. This latter star has however a relatively short observational time coverage, three days.

These results strongly suggest that the stars with the lower $\Delta_{\text{F275W},\text{F814W}}$ values on the ChM, specifically those with $\Delta_{\text{F275W},\text{F814W}} \lesssim -0.18$ in NGC 3201, are binaries. It is interesting to note that barium and CH-type stars form generally via wind mass transfer in fairly wide binaries (Jorissen et al. 1998). Such wide binaries are not likely to survive in a GC because of dynamical interactions, which suggests that the two s -rich stars either may not be in binaries systems today, or that formed via RLOF which can lead to tighter, shorter period binaries more likely to survive in a GC.

The observed distribution in $\Delta_{\text{F275W},\text{F814W}}$ is consistent with simulations of non-interacting binaries (presented in Section 6.1) formed by two stars belonging to the 1G population, which predict these objects on lower $\Delta_{\text{F275W},\text{F814W}}$ values. The fact that 2G stars in the ChM do not show a large spread in $\Delta_{\text{F275W},\text{F814W}}$, as the 1G does, might support the idea of a predominance

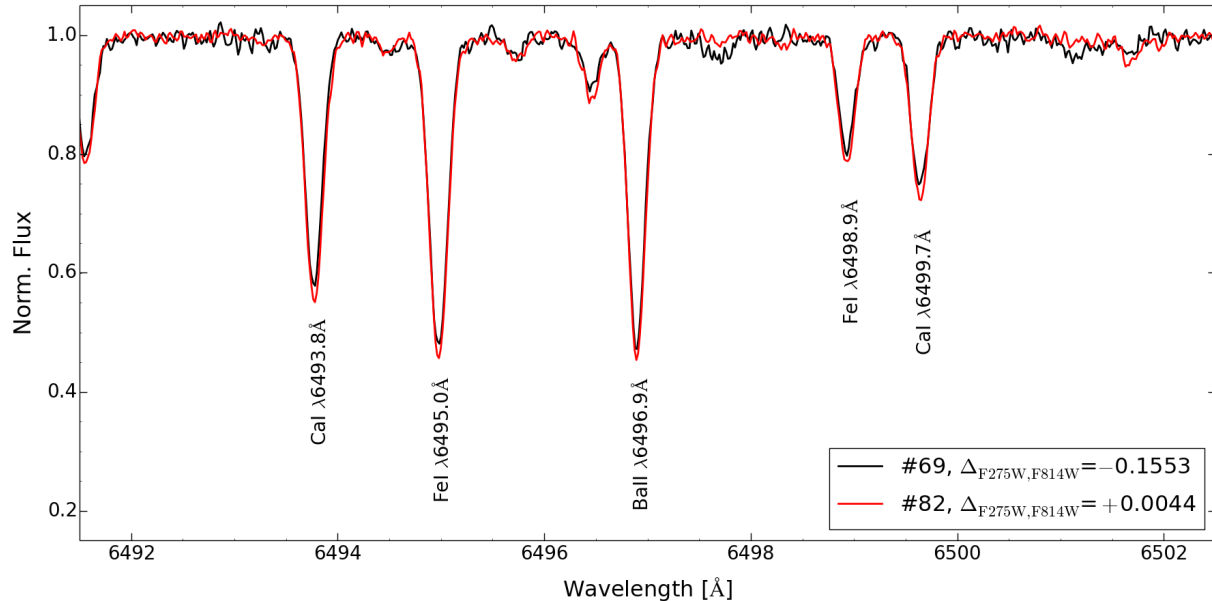


Figure 12. Examples of two portions of spectra including some Fe and Ca analysed spectral features, plus the Ba line $\lambda 6496.9\text{\AA}$. The two represented spectra are for two stars with similar atmospheric parameters, namely #69 ($(T_{\text{eff}}, \log g, [A/H], \xi_t) = (4850 \text{ K}/2.00/-1.55/1.48 \text{ km s}^{-1})$), and #82 ($(T_{\text{eff}}, \log g, [A/H], \xi_t) = (4860 \text{ K}/2.05/-1.44/1.49 \text{ km s}^{-1})$), but different $\Delta_{F275W, F814W}$ values. Overall, the spectral features of #82, with higher $\Delta_{F275W, F814W}$, are deeper.

of 1G-1G binaries in GCs (Lucatello et al. 2015). Furthermore, we note that the BS population accounts for the bluest stars in the elongated 1G population observed on the ChM.

Excluding the three binary candidates, we found a small variation in the overall metallicity in the remaining fifteen giants. $[\text{Fe}/\text{H}]$ has a range of the order of ~ 0.1 dex, which is correlated to the $\Delta_{F275W, F814W}$ value. The absolute abundances of the other analysed elements follow the same variation of Fe, keeping constant the abundance ratios relative to Fe. We exclude the possibility that this small variation in metals is introduced by a change in helium (Y) and a consequent change in Z/X , as in this case we would expect the opposite trend with the $\Delta_{F275W, F814W}$ values.

We can interpret the observed small change in the overall metallicity either as a hint of internal variations among 1G stars, or as an artifact of binarity. In the first case, we have to account for an inhomogeneity in the primordial cloud from which the GC formed. In this context we note that, when data of excellent quality are analysed, very small internal variations in metals are also found in open clusters (Liu et al. 2016). On the other hand, our simulated spectra of binaries suggest that a giant-giant pair (not identical) spectrum is consistent with higher temperature and/or higher metallicity, as the one we find. Previous work on the binary fraction of NGC 3201 however does not support such a

high number of binaries as the one required to account for the observed elongated 1G (Milone et al. 2012b).

We conclude that binarity surely contributes to the elongation of the 1G in NGC 3201. On the other hand, only binaries with $q \gtrsim 0.8$ can produce a sizeable shift towards low $\Delta_{F275W, F814W}$ values. As an additional mechanism, a small inhomogeneity in metals can account for some of the spread among 1G stars in the map. Helium variations, while able to theoretically produce the observed elongation along the 1G on the ChM, seem unlikely.

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ESO MOLECFIT tool (Smette et al. 2014; Kausch et al. 2014).

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Table 1. Coordinates, Photometric Information (including the location on the ChM), Radial Velocities, with associated rms from the # (number) exposures and maximum time between observations (Time interval).

ID	RA	DEC	V	$\Delta_{F275W,F814W}$	$\Delta_C F275W,F336W,F438W$	RV	r.m.s.RV	#	Time interval
	J2000	J2000	mag	mag	mag	[km s ⁻¹]	[km s ⁻¹]		
150	10:17:31.515	-46:24:34.66	14.007	-0.1118	0.0240	490.8	0.22	5	2 months
160	10:17:28.083	-46:24:33.71	14.323	-0.1432	0.0328	493.0	0.23	5	2 months
32	10:17:33.746	-46:25:46.50	14.577	-0.1030	0.0005	498.6	0.14	5	2 months
69	10:17:38.395	-46:25:29.28	14.542	-0.1553	0.0117	493.1	0.17	5	2 months
58	10:17:40.771	-46:24:44.87	13.649	0.0363	-0.0014	498.1	0.19	5	2 months
91	10:17:32.111	-46:25:09.95	14.165	-0.1662	0.0367	493.5	0.18	5	2 months
40	10:17:29.962	-46:25:30.57	13.821	-0.0808	0.0038	496.8	0.21	5	2 months
147	10:17:33.391	-46:24:25.34	14.794	0.0783	0.0030	493.6	0.20	12	8 months
149	10:17:32.115	-46:24:24.49	14.633	-0.2749	0.0938	495.2	0.11	11	3 days
14	10:17:41.299	-46:25:43.50	14.674	-0.0827	0.0077	496.0	0.20	12	7 months
158	10:17:28.360	-46:24:03.89	14.936	-0.1326	0.0343	494.0	0.08	11	3 days
15	10:17:41.161	-46:25:33.81	14.684	-0.0326	0.0042	497.7	0.09	11	3 days
27	10:17:36.395	-46:25:49.35	14.732	0.0026	-0.0143	492.5	0.08	11	3 days
33	10:17:33.501	-46:25:30.71	15.397	-0.1231	0.0600	492.0	0.24	23	8 months
67	10:17:39.149	-46:25:19.59	15.419	-0.1994	0.0804	496.4	0.76	23	8 months
82	10:17:35.886	-46:25:21.01	14.685	0.0044	-0.0197	492.6	0.19	12	7 months
93	10:17:31.507	-46:24:58.17	14.574	-0.2174	0.0858	506.8	0.55	12	7 months
98	10:17:30.035	-46:25:00.41	15.350	-0.1796	0.0809	493.9	0.26	23	8 months

Table 2. Adopted Atmospheric Parameters, derived from spectroscopy, and corresponding Fe I and Fe II abundances (with the associated σ and number of spectral lines #). The last two columns list the T_{eff} and $\log g$ values obtained from photometry.

ID	T_{eff}	$\log g$	[A/H]	ξ_t	$\log\epsilon(\text{FeI})$	σ_{FeI}	$\#_{\text{FeI}}$	$\log\epsilon(\text{FeII})$	σ_{FeII}	#	$T_{\text{effphotometry}}$	$\log g_{\text{photometry}}$
	(K)	(cgs)	(dex)	(km s ⁻¹)	dex			dex			(K)	(cgs)
150	4680	1.60	-1.55	1.53	5.95	0.09	128	6.01	0.075	12	4641	1.76
160	4810	2.00	-1.48	1.48	6.02	0.09	125	6.08	0.065	13	4729	1.94
32	4830	2.00	-1.46	1.44	6.04	0.11	131	6.10	0.075	12	4767	2.06
69	4850	2.00	-1.55	1.48	5.95	0.09	119	6.00	0.072	12	4781	2.05
58	4650	1.60	-1.44	1.64	6.06	0.09	133	6.14	0.051	13	4507	1.54
91	4780	1.90	-1.53	1.54	5.97	0.09	126	6.03	0.067	13	4726	1.87
40	4700	1.70	-1.48	1.60	6.01	0.10	132	6.09	0.059	13	4600	1.66
147	4860	2.07	-1.49	1.40	6.01	0.09	125	6.08	0.084	13	4648	2.08
149	4920	2.15	-1.47	1.46	6.03	0.10	122	6.11	0.066	13	4838	2.12
14	4880	2.15	-1.48	1.54	6.02	0.09	129	6.09	0.060	13	4760	2.09
158	4910	2.20	-1.56	1.42	5.94	0.11	117	6.00	0.060	12	4816	2.23
15	4880	2.20	-1.42	1.48	6.08	0.09	125	6.16	0.036	13	4762	2.10
27	4880	2.10	-1.39	1.48	6.11	0.09	122	6.18	0.070	13	4762	2.12
33	4950	2.35	-1.58	1.37	5.92	0.10	110	5.97	0.071	13	4901	2.45
67	4990	2.36	-1.68	1.41	5.82	0.10	111	5.88	0.090	13	4940	2.48
82	4860	2.05	-1.44	1.49	6.06	0.08	128	6.14	0.087	13	4780	2.11
93	4950	2.33	-1.46	1.55	6.03	0.09	124	6.11	0.055	13	4765	2.06
98	5000	2.53	-1.55	1.36	5.95	0.09	112	6.01	0.083	13	4922	2.44

Table 3. Analyzed Chemical Abundances from Li to Sc.

STAR	A(Li)		[O/Fe]		[Na/Fe]		[Mg/Fe]		[Al/Fe]		[Si/Fe]		[Ca/Fe]		[Sc/Fe]							
	LTE	non-LTE	#	σ	LTE	non-LTE	#	σ	#	σ	#	σ	#	σ	#	σ						
150	<0.37 ^a	-	2	0.06	-0.16	0.02	4	0.34	0.01	2	-0.12	0.25	2	0.32	0.12	7	0.28	0.09	20	-0.05	0.14	6
160	<0.46 ^a	-	1	-	-0.13	0.08	4	0.23	0.01	2	-0.09	-	1	0.28	0.09	6	0.24	0.08	20	0.01	0.14	6
32	0.94 ^a	1.05	2	0.20	-0.11	0.06	3	0.26	0.04	2	-0.10	-	1	0.31	0.15	6	0.26	0.09	20	0.03	0.13	5
69	0.74 ^a	0.84	1	-	-0.20	0.02	3	0.27	0.01	2	<0.00	-	-	0.33	0.09	6	0.26	0.09	20	0.02	0.12	6
58	<0.31 ^a	-	2	0.07	-0.13	0.02	4	0.28	0.04	2	-0.13	0.26	2	0.29	0.10	7	0.28	0.08	20	-0.01	0.13	7
91	<0.37 ^a	-	1	-	-0.13	0.15	4	0.28	0.02	2	-0.10	0.20	2	0.32	0.11	7	0.27	0.08	20	0.00	0.10	6
40	<0.27 ^a	-	2	0.02	-0.20	0.08	4	0.28	0.02	2	-0.23	-	1	0.33	0.11	7	0.26	0.09	20	0.01	0.10	6
147	0.91	1.02	1	-	-0.15	0.09	4	0.29	0.02	2	<-0.10	-	-	0.26	0.09	7	0.27	0.10	20	-0.07	0.13	6
149	<0.50 ^a	-	1	-	-0.06	0.13	4	0.27	0.02	2	-0.18	-	1	0.32	0.11	7	0.26	0.08	20	-0.02	0.14	6
14	0.94	1.04	1	-	-0.18	0.04	4	0.27	0.03	2	-0.13	0.20	2	0.30	0.10	5	0.26	0.09	20	0.00	0.14	6
158	0.95	1.04	1	-	-0.15	0.16	4	0.25	0.02	2	-0.13	-	1	0.33	0.12	5	0.26	0.09	20	0.01	0.12	5
15	0.93	1.03	2	0.09	-0.16	0.06	4	0.26	0.02	2	-0.05	0.30	2	0.32	0.09	5	0.28	0.08	20	0.04	0.12	5
27	0.97	1.07	1	-	-0.28	0.07	4	0.20	0.01	2	-0.26	0.19	2	0.28	0.09	6	0.24	0.09	19	-0.03	0.09	6
33	0.80	0.88	1	-	-0.18	0.04	2	0.24	0.01	2	<0.10	-	-	0.30	0.10	6	0.28	0.07	18	-0.03	0.14	6
67	0.87	0.94	1	-	-0.19	0.21	4	0.27	0.01	2	<-0.10	-	-	0.29	0.10	6	0.24	0.08	19	-0.02	0.20	6
82	0.80	0.91	2	0.02	-0.17	0.10	4	0.22	0.01	2	-0.13	0.08	2	0.28	0.09	7	0.25	0.09	20	-0.03	0.14	6
93	<0.44 ^a	-	2	0.02	-0.16	0.12	4	0.28	0.00	2	-0.33	-	1	0.35	0.10	5	0.23	0.09	20	0.05	0.10	5
98	0.90	0.97	1	-	-0.22	0.22	4	0.22	0.00	2	-0.10	0.38	2	0.28	0.08	7	0.25	0.08	19	-0.04	0.09	6
avg.	0.89	0.98	0.57	-	-0.16	0.02	0.26	-0.23	0.01	0.02	-0.15	0.31	0.26	-0.01	0.01	0.01	0.00	0.00	0.01	0.01	0.01	0.01
\pm	0.02	0.02	0.02	-	0.01	0.01	0.01	0.01	0.01	0.02	0.02	0.01	0.01	0.01	0.01	0.01	0.00	0.00	0.01	0.01	0.01	0.01
σ	0.07	0.08	0.07	-	0.05	0.05	0.03	0.05	0.03	0.08	0.08	0.02	0.02	0.02	0.02	0.02	0.02	0.02	0.02	0.02	0.02	0.02

^a Star brighter than the RGB bump.

Table 4. Analyzed Chemical Abundances from Ti to Ni.

STAR	[Ti/Fe]I	σ	#	[Ti/Fe]II	σ	#	[V/Fe]	σ	#	[Cr/Fe]I	σ	#	[Cr/Fe]II	σ	#	[Mn/Fe]	σ	#	[Co/Fe]	σ	#	[Ni/Fe]	σ	#
150	0.19	0.09	20	0.24	0.05	5	-0.10	0.09	14	-0.08	0.09	4	0.10	0.02	2	-0.50	0.10	5	-0.12	0.07	2	-0.06	0.11	23
160	0.25	0.12	20	0.21	0.09	5	-0.06	0.11	13	-0.08	0.11	4	0.10	0.03	2	-0.46	0.07	5	-0.23	0.00	2	-0.03	0.12	23
32	0.19	0.11	20	0.25	0.03	5	-0.09	0.09	14	-0.09	0.10	4	0.05	0.02	2	-0.45	0.09	5	-0.17	0.18	2	-0.04	0.14	24
69	0.25	0.12	18	0.29	0.05	5	-0.02	0.11	9	-0.15	0.07	4	0.16	0.02	2	-0.49	0.08	5	-0.09	0.01	2	-0.04	0.13	20
58	0.24	0.09	22	0.26	0.04	5	-0.03	0.08	14	-0.06	0.08	4	0.15	0.02	2	-0.40	0.05	5	-0.16	0.15	2	0.01	0.12	25
91	0.19	0.12	20	0.24	0.06	5	-0.11	0.12	13	-0.15	0.12	4	0.07	0.02	2	-0.50	0.12	5	-0.10	0.15	2	-0.04	0.11	22
40	0.20	0.08	21	0.26	0.07	5	-0.06	0.08	14	-0.10	0.08	4	0.10	0.06	2	-0.46	0.06	5	-0.07	0.00	2	-0.02	0.14	24
147	0.24	0.09	19	0.27	0.05	5	-0.10	0.07	13	-0.08	0.10	4	0.10	0.12	2	-0.48	0.12	5	-0.23	0.09	2	-0.05	0.12	21
149	0.17	0.08	19	0.17	0.03	5	-0.07	0.11	11	-0.13	0.08	4	0.05	0.01	2	-0.55	0.12	5	-0.15	-	1	-0.07	0.10	21
14	0.23	0.07	22	0.27	0.08	5	-0.02	0.06	9	-0.05	0.11	4	0.18	0.02	2	-0.44	0.06	5	-0.20	-	1	-0.01	0.10	24
158	0.21	0.11	17	0.21	0.03	5	-0.09	0.10	7	-0.21	0.10	3	0.10	0.05	2	-0.58	0.12	5	999	-	-	-0.04	0.12	20
15	0.20	0.06	21	0.25	0.05	5	-0.02	0.09	11	-0.11	0.06	4	0.11	0.04	2	-0.50	0.10	5	-0.10	0.05	2	-0.02	0.12	24
27	0.15	0.08	21	0.19	0.05	5	-0.04	0.09	10	-0.11	0.06	4	-0.01	0.14	2	-0.50	0.08	5	-0.14	0.03	2	-0.04	0.11	24
33	0.16	0.06	15	0.25	0.05	5	-0.04	0.14	4	-0.19	0.07	3	0.05	0.10	2	-0.54	0.13	5	0.04	-	1	-0.07	0.14	17
67	0.21	0.07	14	0.18	0.11	5	-0.07	0.08	6	-0.25	0.12	2	0.17	0.08	2	-0.50	0.14	5	999	-	-	-0.08	0.12	15
82	0.22	0.08	20	0.24	0.08	5	-0.08	0.12	13	-0.09	0.07	4	0.07	0.01	2	-0.47	0.08	5	-0.20	0.07	2	-0.02	0.10	24
93	0.24	0.09	19	0.32	0.07	5	-0.05	0.07	9	-0.05	0.15	4	0.17	0.02	2	-0.47	0.10	5	-0.19	0.19	2	-0.03	0.10	22
98	0.17	0.08	18	0.22	0.06	5	-0.06	0.17	6	-0.15	0.10	4	0.07	0.08	2	-0.49	0.15	5	999	-	-	-0.07	0.10	18
avg.	0.21			0.24			-0.06			-0.12			0.10			-0.49			-0.14			-0.04		
\pm	0.01			0.01			0.01			0.01			0.01			0.01			0.02			0.01		
σ	0.03			0.04			0.03			0.06			0.05			0.04			0.07			0.02		

Table 5. Analyzed Chemical Abundances from Cu to Eu.

STAR	[Cu/Fe]	[Zn/Fe]	[Y/Fe]	σ	#	[Zr/Fe]	[Ba/Fe]	σ	#	[La/Fe]	σ	#	[Pr/Fe]	[Nd/Fe]	σ	#	[Eu/Fe]
150	-0.57	0.13	-0.24	0.07	3	0.04	0.20	0.04	3	0.17	0.09	7	0.23	0.10	0.03	2	0.45
160	-0.55	0.19	-0.18	0.08	3	0.11	0.33	0.05	3	0.27	0.11	5	0.33	0.25	0.02	2	0.60
32	-0.57	0.18	-0.21	0.07	2	0.03	0.32	0.03	3	0.25	0.07	6	0.32	0.28	0.06	2	0.56
69	-0.65	0.14	-0.20	0.12	2	0.04	0.21	0.04	3	0.25	0.05	6	0.33	0.20	0.05	2	0.47
58	-0.34	0.23	-0.24	0.07	3	0.16	0.25	0.02	3	0.24	0.08	7	0.30	0.27	0.08	2	0.56
91	-0.57	0.07	-0.17	0.15	3	0.02	0.21	0.01	3	0.14	0.05	6	0.32	0.13	0.04	2	0.28
40	-0.50	0.13	-0.27	0.04	3	0.13	0.19	0.06	3	0.18	0.06	6	0.18	0.13	0.07	2	0.43
147	-0.55	0.20	-0.27	0.03	3	-0.01	0.27	0.05	3	0.27	0.06	6	0.40	0.17	0.09	2	0.54
149	-0.64	0.01	-0.19	0.04	3	0.15	0.68	0.04	3	0.47	0.07	6	0.45	0.42	0.01	2	0.45
14	-0.52	0.14	-0.15	0.17	3	0.02	0.30	0.07	3	0.29	0.06	6	0.20	0.19	0.05	2	0.55
158	-0.70	-0.03	-0.26	0.01	2	0.09	0.22	0.08	3	0.20	0.13	2	<0.55	0.09	0.00	2	<0.70
15	-0.61	0.11	-0.19	0.10	3	0.20	0.29	0.07	3	0.28	0.09	4	0.34	0.17	0.00	2	0.50
27	-0.62	0.04	-0.32	0.12	3	0.01	0.19	0.04	3	0.17	0.12	6	0.35	0.08	0.04	2	0.45
33	-0.67	0.21	-0.23	0.10	2	0.22	0.23	0.04	3	0.19	0.01	2	<0.70	0.21	0.00	2	0.41
67	-0.75	-0.01	-0.33	0.06	2	-0.05	0.15	0.10	3	0.19	0.03	2	<0.80	0.10	0.01	2	0.45
82	-0.52	0.12	-0.21	0.07	3	0.08	0.30	0.10	3	0.26	0.08	7	0.30	0.19	0.07	2	0.48
93	-0.59	0.12	0.11	0.07	3	0.33	1.16	0.04	3	0.90	0.07	7	0.80	0.91	0.03	2	0.65
98	-0.60	0.03	-0.30	0.07	2	-	0.23	0.08	3	0.26	0.14	2	<0.50	0.10	0.01	2	0.35
avg.	-0.58	0.11	-0.21			0.09	0.32			0.28			0.35	0.22			0.48
\pm	0.02	0.02	0.02			0.02	0.06			0.04			0.04	0.05			0.02
σ	0.09	0.08	0.10			0.10	0.24			0.17			0.15	0.19			0.09

Table 6. Sensitivity of the derived abundances to the uncertainties in atmospheric parameters ($T_{\text{eff}}/\log g/\xi_t/[A/H]=\pm 50 \text{ K}/\pm 0.20/\pm 0.20 \text{ km s}^{-1}/\pm 0.10$), and uncertainties due to the errors in the EWs measurements or in the χ -square fitting procedure. For reference, we also list the variations due to a change in T_{eff} by $\pm 100 \text{ K}$. We reported the total internal uncertainty (σ_{total}) obtained by the quadratic sum of all the contributors to the error.

	ΔT_{eff}	ΔT_{eff}	$\Delta \log g$	$\Delta \xi_t$	$\Delta [A/H]$	$\sigma_{\text{EWs/fit}}$	σ_{total}
	$\pm 100 \text{ K}$	$\pm 50 \text{ K}$	± 0.20	$\pm 0.20 \text{ km s}^{-1}$	$\pm 0.10 \text{ dex}$		
A(Li)	± 0.15	± 0.10	± 0.02	± 0.03	± 0.00	± 0.10	± 0.15
[O/Fe]	± 0.03	± 0.20	± 0.08	± 0.00	± 0.03	± 0.01	± 0.22
[Na/Fe]	∓ 0.06	± 0.05	∓ 0.02	± 0.04	∓ 0.01	∓ 0.02	± 0.07
[Mg/Fe]	∓ 0.06	± 0.02	∓ 0.04	∓ 0.03	∓ 0.02	∓ 0.02	± 0.08
[Al/Fe]	± 0.05	± 0.27	∓ 0.01	± 0.00	∓ 0.01	± 0.02	± 0.27
[Si/Fe]	∓ 0.08	± 0.02	± 0.03	± 0.05	± 0.03	∓ 0.05	± 0.08
[Ca/Fe]	∓ 0.03	± 0.01	∓ 0.01	∓ 0.01	± 0.00	∓ 0.01	± 0.02
[Sc/Fe] II	± 0.09	± 0.02	± 0.04	± 0.07	± 0.05	± 0.03	± 0.10
[Ti/Fe] I	± 0.03	± 0.01	∓ 0.01	± 0.03	± 0.00	± 0.02	± 0.04
[Ti/Fe] II	± 0.08	± 0.02	± 0.04	± 0.05	± 0.05	± 0.02	± 0.09
[V/Fe]	± 0.07	± 0.03	± 0.02	± 0.07	± 0.03	± 0.02	± 0.09
[Cr/Fe] I	∓ 0.03	± 0.03	∓ 0.06	∓ 0.04	∓ 0.05	∓ 0.01	± 0.092
[Cr/Fe] II	± 0.02	± 0.05	± 0.01	± 0.04	± 0.01	± 0.00	± 0.07
[Mn/Fe]	± 0.03	± 0.07	∓ 0.02	± 0.04	∓ 0.01	± 0.01	± 0.08
[Fe/H] I	± 0.13	± 0.00	± 0.01	∓ 0.04	± 0.01	± 0.06	± 0.07
[Fe/H] II	∓ 0.07	± 0.01	± 0.08	∓ 0.07	± 0.01	∓ 0.02	± 0.11
[Co/Fe]	± 0.04	± 0.06	± 0.04	± 0.09	± 0.05	± 0.00	± 0.13
[Ni/Fe]	∓ 0.01	± 0.01	± 0.01	± 0.04	± 0.01	∓ 0.01	± 0.04
[Cu/Fe]	± 0.14	± 0.10	∓ 0.00	∓ 0.03	∓ 0.01	± 0.07	± 0.13
[Zn/Fe]	∓ 0.14	± 0.05	± 0.05	∓ 0.04	± 0.01	∓ 0.07	± 0.11
[Y/Fe] II	± 0.07	± 0.05	∓ 0.01	± 0.01	± 0.01	± 0.03	± 0.06
[Zr/Fe] II	∓ 0.04	± 0.15	± 0.10	∓ 0.00	± 0.05	∓ 0.02	± 0.19
[Ba/Fe] II	± 0.03	± 0.10	± 0.07	∓ 0.15	± 0.03	± 0.01	± 0.20
[La/Fe] II	± 0.03	± 0.07	± 0.07	∓ 0.01	± 0.03	± 0.01	± 0.10
[Pr/Fe] II	± 0.04	± 0.15	± 0.07	± 0.00	± 0.05	± 0.01	± 0.17
[Nd/Fe] II	± 0.14	± 0.06	± 0.06	± 0.09	± 0.07	± 0.05	± 0.15
[Eu/Fe] II	± 0.00	± 0.15	± 0.08	± 0.00	± 0.04	∓ 0.00	± 0.17

Table 7. Spearman correlation coefficient (r) for each $\log\epsilon(X)$ abundance as a function of $\Delta_{F275W,F814W}$, and probability that the corresponding slope is due to randomness (f).

Element	r	f (%)
Fe	0.68	2.2
Li _{non-LTE}	0.35	27.6
O	0.38	33.1
Na _{non-LTE}	0.56	3.1
Mg	0.70	3.7
Al	0.37	37.6
Si	0.51	9.9
Ca	0.75	0.0
Sc II	0.47	8.2
Ti I	0.68	0.0
Ti II	0.75	2.2
V	0.57	3.5
Cr I	0.68	0.3
Cr II	0.61	0.5
Mn	0.63	0.2
Co	0.09	40.0
Ni	0.67	0.0
Cu	0.66	1.0
Zn	0.65	1.0
Y II	0.37	4.8
Zr II	0.38	25.2
Ba II	0.55	18.7
La II	0.60	2.9
Pr II	0.57	15.5
Nd II	0.52	10.1
Eu II	0.63	4.1

Table 8. Atomic data and equivalent widths for program stars

Wavelength [\AA]	Species	E.P. [eV]	$\log(gf)$	EW [m\AA]
150				
6806.860	26.0	2.730	-3.140	21.9
4885.430	26.0	3.880	-1.150	46.9
4917.240	26.0	4.190	-1.270	32.5
6581.220	26.0	1.490	-4.680	14.7
6592.910	26.0	2.730	-1.490	99.2
6593.880	26.0	2.430	-2.420	76.6
6608.040	26.0	2.280	-3.960	10.8
6609.120	26.0	2.560	-2.690	52.9
6627.560	26.0	4.550	-1.500	6.6

Notes: Only a portion of this table is shown here to demonstrate its form and content. A machine-readable version of the full table will be available.