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MeerKAT-64 discovers wide-spread tidal debris in the nearby NGC 7232 galaxy group

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ABSTRACT

We report the discovery of large amounts of previously undetected cold neutral atomic hydrogen (HI) around the core triplet galaxies in the nearby NGC 7232 galaxy group with MeerKAT. With a physical resolution of ~ 1 kpc, we detect a complex web of low-surface-brightness HI emission down to a 4σ column density level of $\sim 1 \times 10^{19} \text{ cm}^{-2}$ (over 44 km s^{-1}). The newly discovered HI streams extend over ~ 20 arcmin corresponding to 140 kpc in projection. This is approximately three times the HI extent of the galaxy triplet (52 kpc). The HI debris has an HI mass of $\sim 6.6 \times 10^9 M_{\odot}$, more than 50 per cent of the total HI mass of the triplet. Within the galaxy triplet, NGC 7233 and NGC 7232 have lost a significant amount of HI while NGC 7232B appears to have an excess of HI. The HI deficiency in NGC 7232 and NGC 7233 indicates that galaxy–galaxy interaction in the group concentrates on this galaxy pair while the other disc galaxies have visited them over time. In comparison to the AMIGA sample of isolated galaxies, we find that with regards to its total HI mass the NGC 7232/3 galaxy triplet is not HI-deficient. Despite the many interactions associated to the triplet galaxies, no HI seems to have been lost from the group (yet).

Key words: techniques: interferometric – galaxies: groups: individual: NGC 7232 – galaxies: interactions – intergalactic medium – radio lines: galaxies.

1 INTRODUCTION

It is widely known that galaxy properties such as morphology, gas content, and star formation are affected by interactions between galaxies and their environments (Dressler 1980). As the majority of galaxies in the nearby universe reside in small galaxy groups (e.g. Robotham et al. 2011), it is important to understand how galaxies evolve in such environments (Tully 1987; Balogh et al. 2004;

Knobel et al. 2015; Saulder et al. 2016). Observational signatures of interaction histories are seen in the form of, e.g. the length and shape of tidal tails, bridges, and other debris (e.g. Koribalski & Dickey 2004; English et al. 2010; Saponara et al. 2018; Koribalski 2020).

The role of galaxy groups has long been prescribed for understanding galaxy properties over a broad range of environments. In the hierarchical scenario of large-scale structure formation, galaxy groups and the galaxies therein are the building blocks of rich clusters (Kraft et al. 2006). There is evidence that galaxies must undergo some amount of pre-processing (e.g. the change in their gas content and morphology) through tidal interaction before falling into

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clusters (e.g. Zabludoff & Mulchaey 1998; Porter et al. 2008; Seth & Raychaudhury 2020). This indicates that some of the properties of the galaxies that we see in galaxy clusters today are the results of transformation processes pre-dating the infall into clusters and hence taking place under very different conditions (Dressler 2004). As a consequence, galaxy groups are being observed to probe how and by which mechanisms galaxies are pre-processed before they arrive in clusters.

Neutral atomic hydrogen (HI) is known to be an ideal tracer of current as well as past galaxy interactions. This is because the extended HI envelope, which is typically a factor of 2–3 larger than the observed bright stellar body (Pisano, Wilcots & Elmegreen 2000; Bigiel & Blitz 2012; Koribalski et al. 2018) is more susceptible than the optical disc to external disturbances via gravitational and hydrodynamical interactions (Gunn & Gott 1972; Barnes et al. 2001; Lee-Waddell et al. 2016; Spindler et al. 2018).

The effect of the environment on the HI content of galaxies is well studied in dense environments. It is well known that in high-density environment a substantial amount of gas goes into the inter-cluster medium (ICM), making the spiral galaxies in rich clusters to be HI-deficient, i.e. contain less HI mass than field spirals of similar HI morphology and size (e.g. Schröder, Drinkwater & Richter 2001; Solanes et al. 2001; Chung et al. 2009). However, HI-deficient galaxies have also been found in compact groups (e.g. Verdes-Montenegro et al. 2001; Hess et al. 2017; Jones et al. 2019). Over the past few years, HI-deficient galaxies have also been reported in loose groups (e.g. Kilborn et al. 2005; Sengupta & Balasubramanyam 2006; Sengupta, Balasubramanyam & Dwarakanath 2007; Hess & Wilcots 2013). Of particular interest are hence high-resolution studies of the neutral gas in galaxy groups, to study in detail the transformation processes that lead to its removal from galaxies in groups. Such observations used to be difficult in the past time, often requiring a significant amount of observing time, owing to the limited FOV and sensitivity of suitable telescopes. The number of studies of the resolved properties of HI in groups is therefore limited. Due to these restrictions, one particular question that is poorly addressed in the literature is whether and in which form low-column-density gas in groups exists and what its likely fate is.

Previous surveys of the HI content in nearby groups have often revealed the presence of intra-group material (e.g. Ryder et al. 2001; Bekki et al. 2005a; Bekki, Koribalski & Kilborn 2005b). These observations were typically conducted with large single-dish telescopes (e.g. HI Parkes All Sky Survey, HIPASS, Barnes et al. 2001, and ALFALFA, Haynes et al. 2018). While single-dish telescopes provide good sensitivity for observing a wide area of the sky in a reasonable amount of time, the coarse spatial resolution of these instruments does not allow to explore the low-column-density universe on small spatial scales, thus the need for complementary deep, high-angular-resolution HI observations. MeerKAT with its high-sensitivity and high-spatial-resolution capabilities enables us to better understand how the low-column-density phase of HI on small spatial scales relates to galaxy evolution (e.g. Kleiner et al. 2021).

The nearby NGC 7232 galaxy group is located at a distance of ~ 24 Mpc (Garcia 1995), which we adopt throughout this paper. This field displays a diverse array of gaseous features, apparent galaxy interactions, as well as gas-rich and gas-poor galaxies, thus making it an excellent target to study how such interactions may impact galaxy evolution. Fig. 1 shows a *grz* optical image of the NGC 7232 galaxy group obtained from the Dark Energy Legacy Survey¹ (Dey

et al. 2019). Using the HIPASS, we estimate that the group has 19 members (Koribalski et al. 2004), 4 of which are bright galaxies. The three spirals, NGC 7232, NGC 7232B, and NGC 7233, lying within 5 arcmin in projection (~ 35 kpc) of each other make up the NGC 7232/3 triplet, while the lenticular galaxy IC 5181 lies 10 arcmin south-west of the triplet.

The first interferometric HI studies of the NGC 7232 galaxy group (Barnes & Webster 2001) were carried out with the Australia Telescope Compact Array (ATCA). The ATCA HI intensity distribution revealed traces of recent interactions within the group, including HI streams connecting the triplet galaxies. Single-dish HIPASS observations of the NGC 7232 galaxy group show the triplet galaxies are embedded in a common HI envelope (Koribalski & Dickey 2004). However, the Parkes observations do not have the spatial resolution to separate these galaxies. Neighbouring galaxies around the galaxy group were also reported. Recently, Lee-Waddell et al. (2019) obtained and analysed ASKAP-12 Early Science data of the NGC 7232 galaxy group. These observations were conducted as part of the Wide-field ASKAP L-Band Legacy All-sky Blind survey (WALLABY; Koribalski et al. 2020) Early Science program. From these observations, 17 HI sources were reported, 5 of which were new galaxies while the other 6 detections appear to be tidal debris in the form of HI clouds that are associated with the triplet NGC 7232/3 and the S0 galaxy IC 5181.

While NGC 7232 is a well-studied group, neither of the previous two interferometric HI studies (Barnes & Webster 2001; Lee-Waddell et al. 2019) were able to recover all the HI emission detected by HIPASS (Koribalski et al. 2004), leaving an open question as to where the missing gas resides. With column density sensitivities of $\sim 10^{20}$ cm⁻², recent ASKAP observations (Lee-Waddell et al. 2019) indicated the presence of undetected HI gas that could form a bridge connecting the galaxy triplet NGC 7232/3 to the surrounding HI clouds. It is likely that an increase in sensitivity could enable the detection of more low-column-density HI gas and provide an understanding to where the gas resides within the galaxy triplet. This motivated our deeper MeerKAT observations. We observed the nearby NGC 7232 galaxy group at 1.4 GHz with MeerKAT using the 4k mode (44 km s⁻¹). At comparable velocity resolution, the sensitivity of MeerKAT allows us to reach a five times lower rms compared to the published ASKAP data (Lee-Waddell et al. 2019). Although the coarse velocity resolution of our observations does not allow us to carry out a detailed analysis of the HI in and around the galaxy triplet, the sensitivity of MeerKAT spectral line observations has allowed us to detect spectacular new HI debris at column densities of $\sim 10^{19}$ cm⁻². The main focus of this paper is to study the HI distribution of the newly detected wide-spread HI debris around the NGC 7232/3 triplet (shown in the white circle, Fig. 1) with the aim of understanding the interaction processes between the galaxy triplet with its surrounding environment.

This paper is structured as follows. In Section 2, we describe the MeerKAT observations and data reduction. Section 3 presents the results of the NGC 7232/3 galaxy triplet and its surroundings. Sections 4 and 5 present our discussion and summary, respectively.

2 MEERKAT OBSERVATIONS AND DATA REDUCTION

The HI observations of the galaxy group NGC 7232 were obtained with the MeerKAT telescope (Camilo et al. 2018) between 2019 May and June. The data were obtained in three epochs using a single pointing centred on the NGC 7232/3 triplet (α , δ (J2000) = 22:15:38.4, $-45:51:00.3$). A total of 61 antennas were used, providing

¹www.legacysurvey.org.

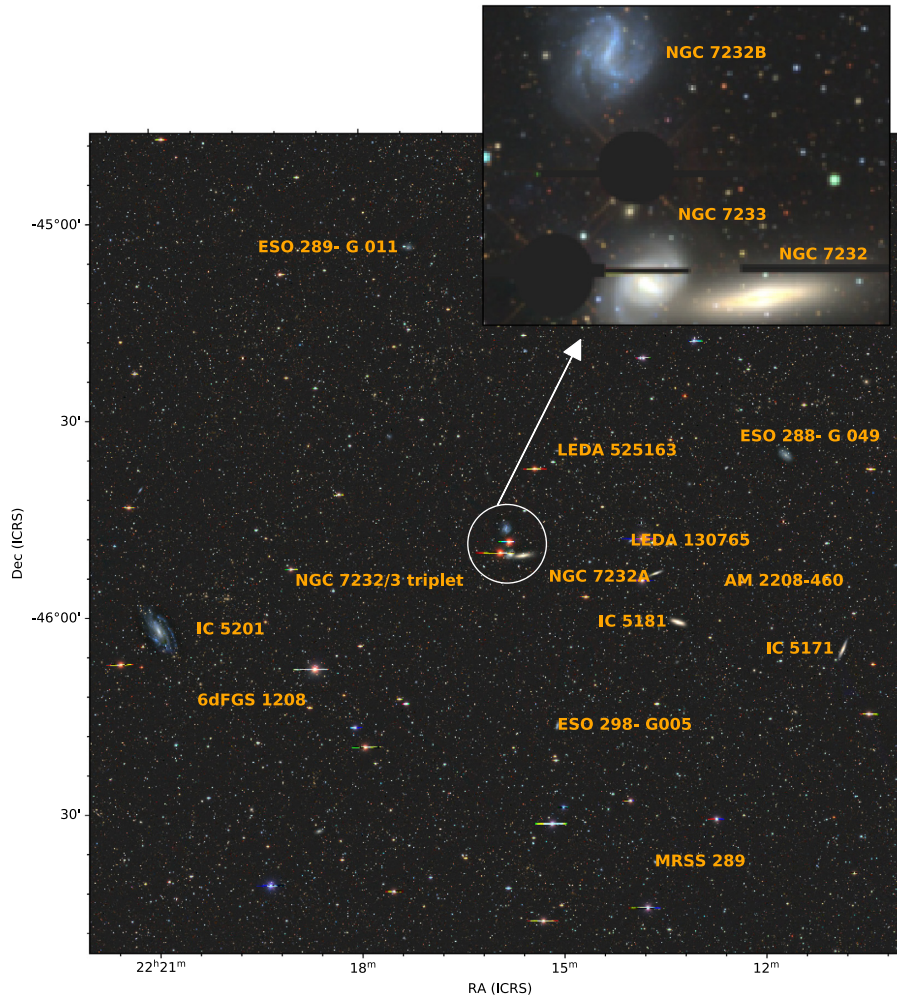


Figure 1. Three-colour (*grz*) optical image of the NGC 7232 group from the Dark Energy Legacy Survey with the MeerKAT H I-detected galaxies labelled. The white circle indicates the NGC 7232/3 galaxy triplet. The top panel shows the NGC 7232/3 galaxy triplet with bright foreground stars and related artefacts obscured.

an excellent *uv* coverage. The observations were conducted in the *L* band (900–1670 MHz) using the SKARAB correlator in 4k-wideband mode (8-s integration time), centred at a frequency of 1283.8 MHz. In this mode, the band is divided into 4096 channels with a channel width of 209 kHz ($\sim 44 \text{ km s}^{-1}$ for H I at $z = 0$). We observed the bright radio source PKS B1934–6342 as the standard flux and bandpass calibrator for 10 min every ~ 2 h and PKS 2214–3835 as a phase calibrator for 2 min every ~ 15 min. The total integration time on source was 15 h. Table 1 summarizes the parameters of the MeerKAT observations.

The data were reduced using the CARACAL² pipeline (Józsa et al. 2020; Józsa et al. 2020; Makhathini et al., in preparation). CARACAL is an automated H I and radio continuum data reduction and analysis pipeline. The pipeline runs several open-source radio interferometry software packages in containerized environments provided by STIMELA,³ a PYTHON-based scripting framework. All data reduction software described below was used in a containerized version inside this framework. We processed a bandwidth of 30 MHz, covering a frequency range between 1400 and 1430 MHz (recession

Table 1. Parameters of the MeerKAT observations.

| Parameter | NGC 7232 group |
|--------------------------------------------------------------------------------------|----------------------------------------------------------------------------------|
| Start of observations | 2019 May |
| End of observations | 2019 Jul |
| Number of antennas | 60–64 |
| Baseline range | 29 m–8 km |
| Total integration | 15 h on source |
| FWHM of primary beam | $\sim 1^\circ$ |
| Channel width | 209 kHz (44 km s^{-1}) |
| Number of channels | 4096 |
| Flux/bandpass calibrator | PKS B1934–6342 |
| Phase calibrator | PKS 2214–3835 |
| Robust = 0 weighting function tapered to 30 arcsec | |
| FWHM of synthesized beam | $41.9 \times 35.5 \text{ arcsec}^2$ ($4.8 \text{ kpc} \times 4.1 \text{ kpc}$) |
| rms noise | 0.1 mJy |
| H I column density limit (4σ over 1 channel = 44 km s^{-1}) | $\sim 1 \times 10^{19} \text{ cm}^{-2}$ |

²<https://caracal.readthedocs.io/en/latest/>.

³<https://github.com/ratt-ru/Stimela/wiki>.

velocities between 4244 and -1930 km s^{-1}), centred around the H I line of the NGC 7232 group.

We first flagged bad data due to radio frequency interference (RFI) with AOFLAGGER (Offringa & Smirnov 2017). For calibrators, channels with Galactic H I emission were discarded (1417–1421 MHz). The corrections to the flux, delay terms, bandpass shapes and antenna gains were determined using the CASA (McMullin et al. 2007) task *bandpass* while the time-varying phases and antenna gains were corrected using the CASA task *gaincal*. Calibration solutions were then applied to the target and the corrected target visibilities were split from the calibrators using the CASA task *mstransform*. The corrected target visibilities were then flagged with AOFLAGGER. To improve the data quality, phase-only antenna gain corrections were derived with CUBICAL (Kenyon et al. 2018), in a self-calibration loop in combination with WSCLEAN (Offringa et al. 2014). Solutions were derived for every 128 s of data.

Imaging of the calibrated data was performed using WSCLEAN. Before spectral line imaging, we replaced the corrected data column with the difference between the corrected data and the model data. This was to subtract the continuum clean model generated from the self-calibration process. An additional continuum subtraction was performed by fitting a third-order polynomial to line-free channels using the *uvlin* option in the CASA task *mstransform*. Having separated the H I emission from the continuum, a data cube was generated using a Robust weighting of 0. For the deconvolution, we used a threshold of 0.5σ within CLEAN regions created with the Source Finding Application (SOFIA; Serra et al. 2015). The restoring beam of the cube was $10.4 \times 7.2 \text{ arcsec}^2$. To explore the low-column-density H I emission, a Gaussian taper of 30 arcsec was applied to produce a second cube with a synthesized beam of $41.9 \times 35.5 \text{ arcsec}^2$ ($4.8 \text{ kpc} \times 4.1 \text{ kpc}$) and rms noise of $0.1 \text{ mJy beam}^{-1}$. This corresponds to a 4σ column density sensitivity of $1 \times 10^{19} \text{ cm}^{-2}$ over the velocity range of 44 km s^{-1} (one channel). For the analysis presented in this paper, we use the cube produced with a Gaussian taper of 30 arcsec. For comparison, WALLABY reported an rms noise of $\sim 1.6 \text{ mJy beam}^{-1}$ per 4 km s^{-1} channel (or $0.5 \text{ mJy beam}^{-1}$ for 44 km s^{-1}) and 30-arcsec resolution (Koribalski 2012; Koribalski et al. 2020).

The H I intensity and velocity field maps were created using the 3D source finding application, SOFIA. The smooth and clip method with a signal-to-noise cutoff of 4σ was applied. We allowed SOFIA to independently estimate the noise in each channel by fitting a half-Gaussian distribution to the negative pixels. The reliability threshold was set to 100 per cent. We used two spatial smoothing (Gaussian) kernels of one and two times the beamwidth over one channel. No smoothing was applied in velocity space because of the low velocity resolution (44 km s^{-1}) of the data. We checked the setup by independently creating source masks and moment maps based on a 4σ cutoff in each channel employing CASA and MIRIAD, with a similar outcome.

3 RESULTS

3.1 H I distribution of the NGC 7232/3 triplet

The main result from our observations is the detection of wide-spread H I emission around the galaxy triplet. Only the bright H I cloud complex between the NGC 7232/3 triplet and IC 5181 was formerly catalogued as HIPASS J2214–45 (Koribalski et al. 2004). Fig. 2 shows the H I channel maps of the galaxy triplet over the velocity range 1606–2232 km s^{-1} . While the channel width 44 km s^{-1} is

coarse, it allows us to trace the different features associated with the NGC 7232/3 triplet. H I detections already discussed in Koribalski et al. (2004), Barnes & Webster (2001), and Lee-Waddell et al. (2019) are shown in blue and red, while all the unlabelled H I emission shown at different velocities is detected for the first time, with MeerKAT-64. The H I distribution around the galaxy triplet is complex. Apart from the low-column-density H I detected around the galaxy triplet in all channels, we see a prominent H I tail between velocities of 1829–1919 km s^{-1} . The H I tail stretches towards the lenticular galaxy IC 5181, which itself is not detected in H I. The H I tail has a length of $\sim 13.5 \text{ arcmin}$ (94 kpc at 24 Mpc) in projection. We confirm six H I clouds (C1–C6) that were first detected with ASKAP (Lee-Waddell et al. 2019).

Fig. 3 shows the column density map of the galaxy triplet and its surroundings. The bottom left-hand inset panel shows H I emission at the column density sensitivity of the ASKAP data $\geq 10^{20} \text{ cm}^{-2}$ while the top inset panel shows the newly detected H I debris with MeerKAT detected at column densities of $\leq 10^{20} \text{ cm}^{-2}$. Fig. 4 shows an interactive 3D representation of the galaxy triplet and its surrounding environment created using the X3D pathway introduced in Vogt et al. (2016). The 3D rendering highlights different components of the system and how these components are connected in both position and velocity. The green colour represent the low-column-density H I debris while the blue colour shows intermediate column density H I structures. The high-density H I features are shown in red. Using the primary-beam corrected H I intensity map, we measure a total flux of $F_{\text{HI}} = (51.01 \pm 0.29) \text{ Jy km s}^{-1}$ and derive a total H I mass of $(6.91 \pm 0.03) \times 10^9 M_{\odot}$ for the triplet and its surrounding tidal debris. The face-on spiral galaxy NGC 7232B, which is the northern member of the NGC 7232/3 triplet, has a total H I flux of $F_{\text{HI}} = (9.45 \pm 0.31) \text{ Jy km s}^{-1}$ and an H I mass of $(1.28 \pm 0.04) \times 10^9 M_{\odot}$. We estimate H I fluxes of (4.71 ± 0.26) and $(3.32 \pm 0.34) \text{ Jy km s}^{-1}$ for the spiral galaxies NGC 7232 and NGC 7233. This corresponds to H I masses of $(6.40 \pm 0.35) \times 10^8$ and $(4.51 \pm 0.46) \times 10^8 M_{\odot}$, respectively. That makes a total flux of $F_{\text{HI}} = (17.48 \pm 0.51) \text{ Jy km s}^{-1}$ (H I mass of $(2.31 \pm 0.06) \times 10^9 M_{\odot}$) for the galaxies in the triplet and $F_{\text{HI}} = (33.53 \pm 0.19) \text{ Jy km s}^{-1}$ (H I mass of $(4.54 \pm 0.03) \times 10^9 M_{\odot}$) for the tidal H I debris detected with MeerKAT. The newly discovered H I debris is located around the two H I clouds C1 and C2 as identified by Lee-Waddell et al. (2019) with each cloud having a total flux of $F_{\text{HI}} = (1.60 \pm 0.09)$ and $(4.41 \pm 0.05) \text{ Jy km s}^{-1}$. This corresponds to H I masses of $(2.17 \pm 0.01) \times 10^8$ and $(5.90 \pm 0.06) \times 10^8 M_{\odot}$, respectively. Lee-Waddell et al. (2019) report only 7 Jy km s^{-1} associated with the H I debris (clumps C1–6). MeerKAT detects approximately five times more low-column-density H I emission in filaments and clouds than ASKAP. The HIPASS total fluxes for the NGC 7232/3 triplet and its neighboring gas clouds are $F_{\text{HI}} \sim 46.0$ (Meyer et al. 2004) and $\sim 34.6 \pm 4.1 \text{ Jy km s}^{-1}$ (Koribalski et al. 2004). MeerKAT hence recovers all H I detected in Parkes single-dish observations in HIPASS and presumably the total H I content of the NGC 7232 galaxy group (see Fig. 6). Table 2 summarizes the H I properties of the galaxy triplet and its surroundings.

3.2 H I kinematics of the NGC 7232/3 triplet

Although the coarse velocity resolution of our observations does not allow for the detailed analysis of the kinematics of the triplet galaxy group and its surroundings, it is clear from the Fig. 5 that NGC 7232B has higher velocities and is kinematically different from NGC 7232 and NGC 7233 (velocity difference of up to 400 km s^{-1}).

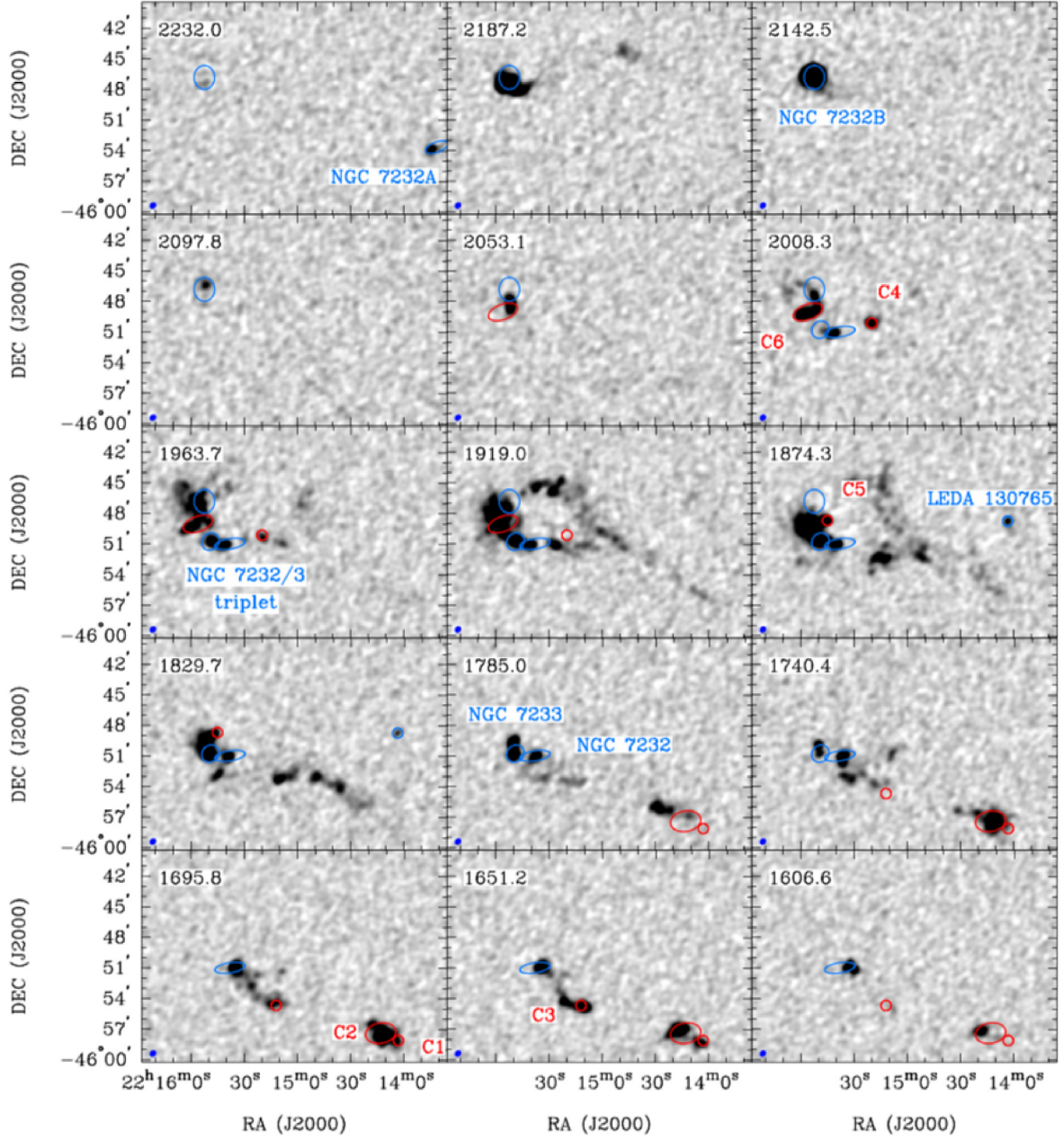


Figure 2. MeerKAT-64 HI channel maps (primary beam corrected) of the NGC 7232/3 triplet, surrounding galaxies, extended tidal streams, and debris. Known galaxies are highlighted by blue symbols and labels, while HI clouds previously detected with ASKAP by (Lee-Waddell et al. 2019) are marked in red. The velocities (in km s^{-1}) are shown in the top left-hand side of the single panels and the synthesized beam ($42.0 \times 35.5 \text{ arcsec}^2$) in the bottom left-hand side.

The HI debris appears to have similar velocities to NGC 7233 and NGC 7232 within $\sim 200 \text{ km s}^{-1}$, which could indicate that most of this gas is coming from the interaction between the galaxy pair and not NGC 7232B.

3.3 HI deficiency equation of the NGC 7232/3 triplet

We determine the HI deficiency of each galaxy in the NGC 7232/3 triplet group to determine if the group spirals are HI-deficient or HI-rich. The HI deficiency is derived by comparing the observed HI mass to the expected HI mass for an isolated galaxy. The expected HI masses ($M_{\text{HI}}^{\text{exp}}$) are derived from the absolute B -magnitudes obtained

from HyperLeda⁴ (Makarov et al. 2014). We have used the scaling relation derived in (Jones et al. 2018) based on measurements or constraints on the HI masses of 844 isolated galaxies compiled by the AMIGA project (Analysis of the interstellar Medium in Isolated Galaxies; Verdes-Montenegro et al. 2005):

$$\log [M_{\text{HI}}^{\text{exp}}/M_{\odot}] = 0.94 \log [L_{\text{B}}/L_{\odot}] + 0.18, \quad (1)$$

where $\log [L_{\text{B}}/L_{\odot}]$ is derived by

$$\log [L_{\text{B}}/L_{\odot}] = 10 + 2 \log [D/\text{Mpc}] + 0.4(M_{\text{bol},\odot} - B_{\text{c}}), \quad (2)$$

⁴<http://leda.univ-lyon1.fr/>.

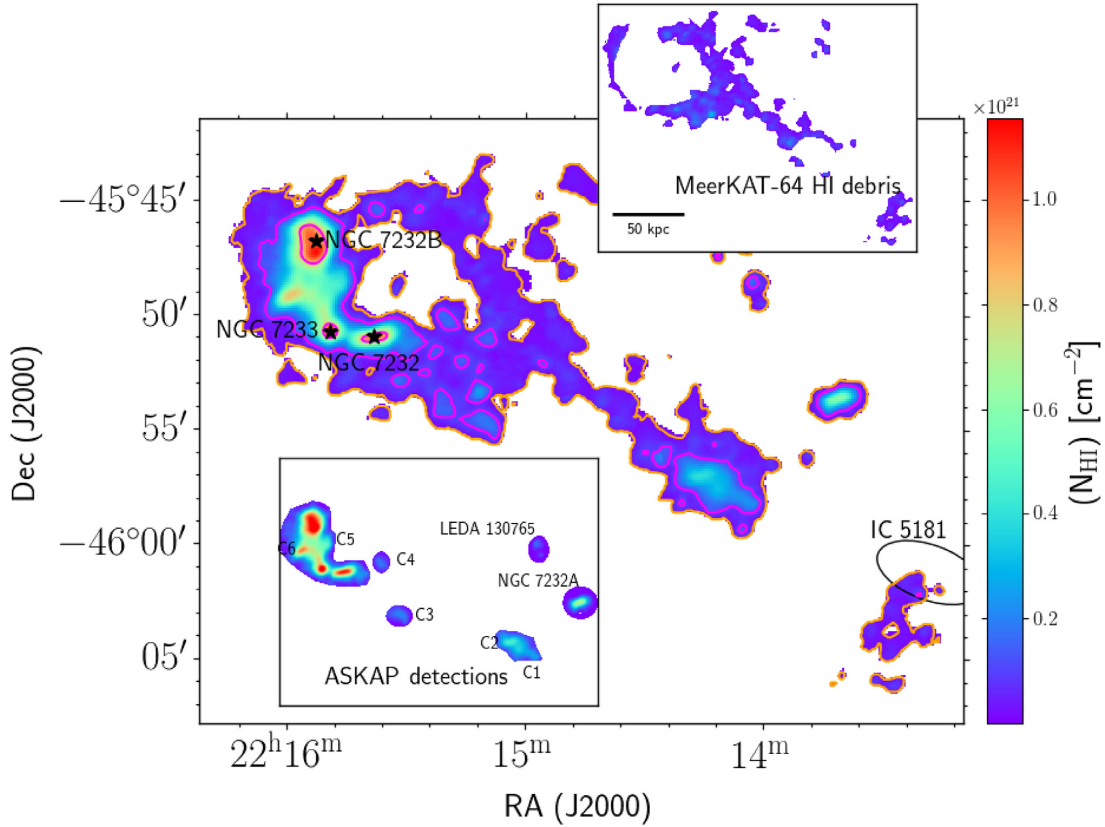


Figure 3. MeerKAT-64 HI column density map (primary beam corrected) of the NGC 7232/3 triplet and its surrounding environment over a velocity range of $\sim 1606\text{--}2232\text{ km s}^{-1}$. The black stars indicate the optical centres of the triplet galaxies. Known galaxies and previously identified HI clouds (Lee-Waddell et al. 2019) are shown in the bottom left-hand inset panel. The lenticular galaxy, IC 5181 ($v_{\text{opt}} = 1987\text{ km s}^{-1}$), marked in an ellipse, is not detected in HI. Main panel: The HI column density contour levels are $4\sigma \times (1, 16, 80)$, where the 4σ column density limit is $1 \times 10^{19}\text{ cm}^{-2}$. The colours represent HI column density levels of $\geq 10^{20}$ (magenta) and $1 \times 10^{19}\text{ cm}^{-2}$ (orange). The MeerKAT detected HI debris are shown in the top right-hand inset panel.

where $M_{\text{bol}, \odot}$ is the Sun’s bolometric absolute magnitude, we adopt $M_{\text{bol}, \odot} = 4.88$ (Lisenfeld et al. 2011), D is the distance in Mpc, and B_c is the absolute B -magnitude corrected for extinction. The HI deficiency is then defined by

$$\text{DEF} = \log [M_{\text{HI}}^{\text{exp}}/M_{\odot}] - \log [M_{\text{HI}}^{\text{obs}}/M_{\odot}], \quad (3)$$

where $M_{\text{HI}}^{\text{obs}}$ is the HI mass derived from the observations. Positive values of DEF indicate HI deficiency while negative values indicate HI excess. We compare the expected HI mass for each galaxy to its HI mass derived from our observations (see Table 2). NGC 7232 and NGC 7233 are HI-deficient while NGC 7232B is found to have excess HI. Taking into account the uncertainties in B -magnitudes and the scatter in the scaling relation, we consider a galaxy to be HI-deficient when $\text{DEF} > 0.2$ and HI excess if $\text{DEF} < -0.2$. The total HI mass of the NGC 7232/3 triplet including all the HI debris (i.e. HI streams and clouds) is $\log[M_{\text{HI}}/M_{\odot}] = 9.61 \pm 0.11] = 9.84 \pm 0.002$. Comparing the total derived HI mass to the expected HI mass of $\log[M_{\text{HI}}/M_{\odot}] = 9.61 \pm 0.11$, we find a total DEF of -0.23 ± 0.11 , which indicates that the NGC 7232/3 galaxy triplet has an overall HI excess compared to field galaxies.

4 DISCUSSION

One of the questions coming from our data is the possible origin of the HI debris observed. In this work, we search for signs of interaction within the NGC 7232/3 galaxy triplet and its surrounding environ-

ment. The detection of wide-spread HI debris around the galaxy triplet demonstrates the need for more sensitive, high-resolution HI observations if we are to have a complete understanding of how the low-column-density universe on small spatial scales contributes to the overall evolution of group galaxies.

4.1 HI content of the NGC 7232/3 triplet

The presence of a large amount of HI around the galaxy triplet, disturbed outer HI discs, and the presence of a bridge connecting the three triplet members are clear evidence, showing that they are currently interacting within the last Gyr (Barnes & Webster 2001; Koribalski et al. 2004; Lee-Waddell et al. 2019). One way of investigating whether galaxies in a group are undergoing transformation is by looking at their relative HI content. We determine the HI deficiency (DEF) of each galaxy. A comparison of the predicted HI masses to the observed HI masses is shown in Table 2. We find that NGC 7232 and NGC 7233 display an HI deficiency while NGC 7232B is in excess. The HI velocity field map (see Fig. 5, bottom panel) shows that NGC 7232B is significantly offset compared to NGC 7232 and NGC 7233. This could suggest that NGC 7232B has not lost a large amount of gas to the intergalactic medium of the triplet, and that most of the extended HI gas probably originated from the galaxy pair. Considering the total HI mass of the galaxy triplet and its surrounding, we find that the entire triplet is not deficient in HI (-0.23 ± 0.11). HI deficiencies or excess is not

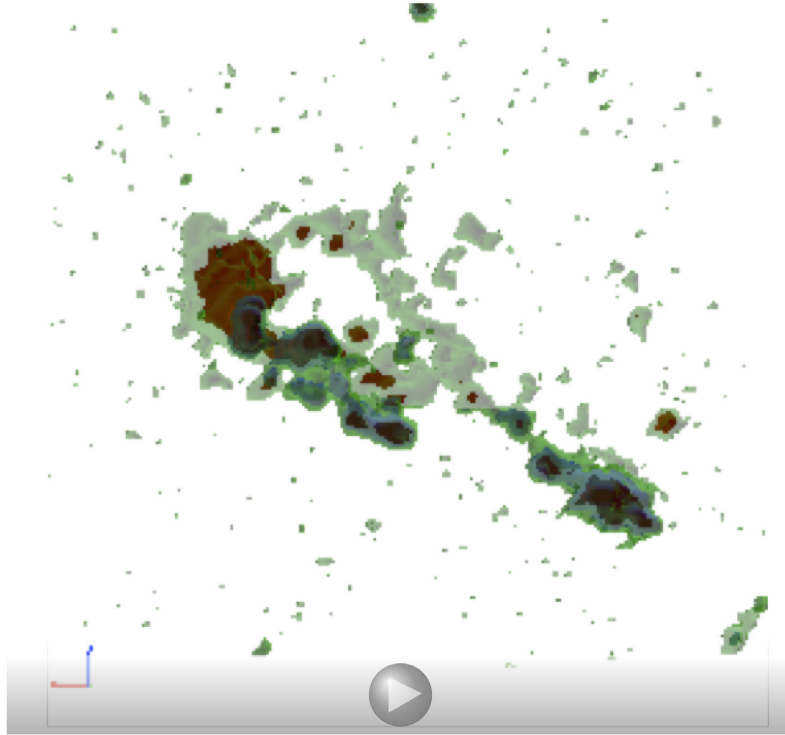


Figure 4. 3D representation of the H I around the galaxy triplet and its surrounding environment. The green components represent the low-column-density H I emission, blue colour shows the intermediate column density structures, and the red components represent the high-column-density regions. This must be viewed using Adobe Acrobat.

Table 2. H I properties of the NGC 7232/3 galaxy triplet and the H I debris.

| Galaxy | F_{HI} (Jy km s $^{-1}$) | v_{HI} (km s $^{-1}$) | $\log [M_{\text{HI}} \text{ obs}/M_{\odot}]$ | $\log [M_{\text{HI}} \text{ pred}/M_{\odot}]$ | H I deficiency |
|------------------|---------------------------------------|------------------------------------|----------------------------------------------|-----------------------------------------------|------------------|
| NGC 7232B | 9.45 ± 0.31 | 2142 ± 17 | 9.11 ± 0.01 | 8.86 ± 0.21 | -0.25 ± 0.21 |
| NGC 7232 | 4.71 ± 0.26 | 1740 ± 120 | 8.80 ± 0.01 | 9.31 ± 0.21 | 0.51 ± 0.21 |
| NGC 7233 | 3.32 ± 0.34 | 1958 ± 28 | 8.65 ± 0.03 | 9.12 ± 0.21 | 0.47 ± 0.21 |
| H I debris | 33.53 ± 0.19 | 1963 ± 82 | 9.65 ± 0.003 | | |
| Triplet + debris | 51.01 ± 0.29 | 1885 ± 115 | 9.84 ± 0.002 | 9.61 ± 0.11 | -0.23 ± 0.11 |

Notes. (1) Galaxy name; (2) integrated H I flux; (3) H I systemic velocity; (4) observed H I mass; (5) predicted H I mass derived using equation (1); and (6) calculated H I deficiencies.

unique to group galaxies (see e.g. Verdes-Montenegro et al. 2001; Kilborn et al. 2005; Jones et al. 2019). Verdes-Montenegro et al. (2001) studied a large sample of compact groups and proposed an evolutionary sequence that such groups follow. In the first phase, the galaxies appear mostly undisturbed in H I; in the second phase, significant quantities of H I appear in the intragroup medium; and in the final stage, the majority of the H I gas is either completely absent or outside the galaxies. Although that study did not consider triplets, the state of the NGC 7232/3 triplet is analogous to a phase 2 compact group. A similar result is seen for HGC 16 (Jones et al. 2019).

4.2 Origin of the H I debris around the NGC 7232/3 galaxy triplet

The unsettled gas around the galaxy triplet strongly indicates interaction between the galaxies and their surrounding environment. The morphology of the newly detected H I debris is complex and extends out to ~ 140 kpc in projection. One prominent feature is an H I tail that extends out to ~ 94 kpc in projection. Given the

length of the H I tail and the mean velocity dispersion (σ) of the triplet galaxy (~ 35 km s $^{-1}$), the age of the H I tail must be at least $94 \text{ kpc}/35 \text{ km s}^{-1} \simeq 2.6$ Gyr. The H I tail and other surrounding H I gas were not detected in the ASKAP observations by Lee-Waddell et al. (2019) (see Fig. 3). The total H I mass contained in the debris ($\sim 4.5 \times 10^9 M_{\odot}$) is more than two times the total H I content of the galaxy triplet (see Fig. 6).

Hydrodynamical simulations have shown that the group's tidal field can be responsible for the stripping of the H I from the outer disc of gas-rich galaxies (Bekki et al. 2005a). The disturbed stellar- and gas morphologies associated with the NGC 7232/3 galaxy triplet is consistent with the scenario that the dominant mechanism in this group is via tidal stripping through galaxy–galaxy interactions. Such tails of H I debris have been observed in numerous galaxy systems (Serra et al. 2012; Saponara et al. 2017; Jones et al. 2019).

The entire environment around the galaxy triplet including gas and stars (see Fig. 5) does not cover a large volume. This could suggest that the system is gravitationally bound to the triplet. The cloud at the south of LEDA 130765 coincides in velocity quite well with the

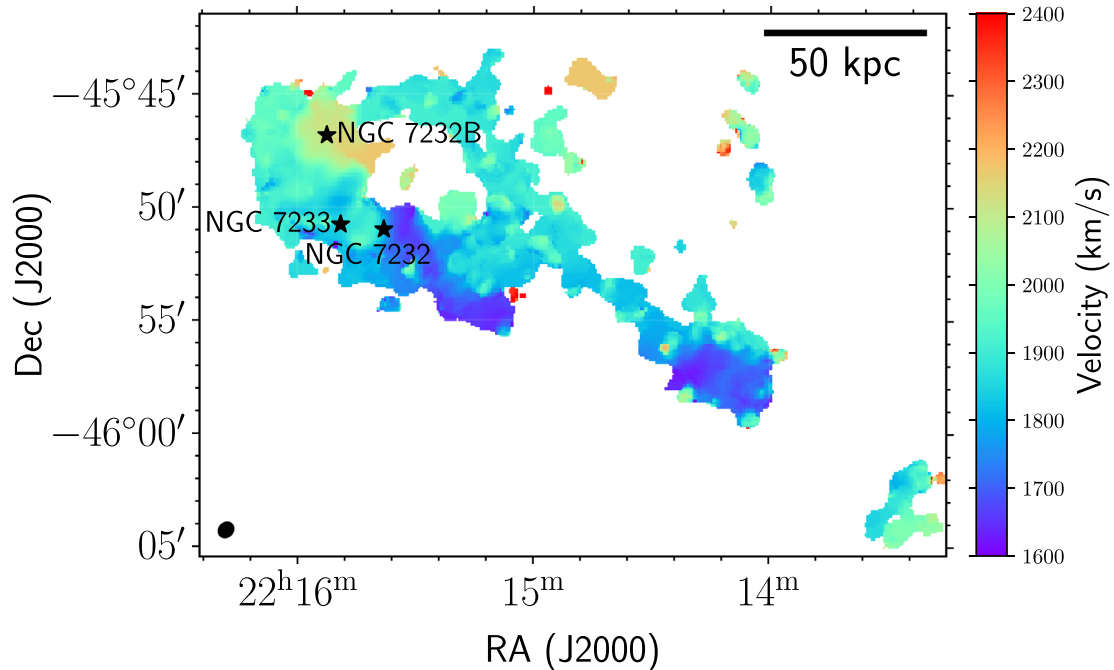


Figure 5. MeerKAT-64 H I velocity field map of the NGC 7232/3 triplet. The black stars indicate the optical centres of the triplet galaxies.

disc of the galaxy. The faded bridge/tail directed towards the triplet could indicate that this galaxy was involved in an earlier interaction with the triplet. Fig. 6 shows the global spectral profile of the galaxy triplet with its surrounding environment.

The profile indicates that most of the H I in the bridge/tail is associated with the NGC 7332/3 galaxy pair and not NGC 7232B. One can also see from the profile that although the H I gas is spatially disturbed it is well aligned in the spectral domain. While NGC 7232B appear to be more in the background based on their velocities (see Table 2), the velocities of the H I features are consistent with that of NGC 7233 and NGC 7232. Thus, it is possible that the H I debris are originating from NGC 7232 and NGC 7233 only.

Fig. 7 shows channel maps of the H I cloud near IC 5181. The H I tail towards IC 5181 is seen around 1874 km s^{-1} . The alignment of the lenticular galaxy with the H I tail suggests that this galaxy could have interacted with the triplet and contributed some of its gas hence speeding up its own evolution.

In short, our observations indicate that there is interaction taking place within the NGC 7232/3 galaxy triplet. However, in the future, it will be interesting to combine H I observations with simulations to determine the origin of the H I debris seen with MeerKAT.

4.3 H I clouds around the NGC 7232/3 triplet

Tidal interactions between gas-rich galaxies can create prominent gaseous tidal features that, in some cases, may form dense clumps of gas and possibly stars. Most of this tidally stripped material will eventually fall back into the interacting galaxies. However, during this process self-gravitating bodies, with comparably higher masses, can be formed out of the tidal debris. Bournaud & Duc (2006) show that as large tidally-formed clouds of H I evolve, if they acquire sufficient H I mass ($> 10^8 M_{\odot}$) and move a sufficient distance away from their parent galaxies, they can decouple from their host tails and become kinematically distinct, long-lived tidal dwarf galaxies

(TDGs). However, without the linking tidal tails, it becomes more difficult to discern a TDG from an accreted satellite dwarf.

We confirm the six H I clumps detected with ASKAP (Lee-Waddell et al. 2019) and report one new faint H I cloud located near the gas-poor lenticular galaxy IC 5181 (see Fig. 3). C3 and C4 are relatively low-mass H I clumps. We measure H I masses of 6.7×10^7 and $9.5 \times 10^7 M_{\odot}$ for C3 and C4, respectively. The low H I masses of C3 and C4 show that these H I clumps are likely short-lived and will probably be reaccreted by the parent spiral galaxies. These two H I clumps are located within the low-column-density complex of H I emission surrounding the galaxy triplet, which makes it difficult to determine their exact origin.

C2 is the most gas-rich H I clump detected near the galaxy triplet with an H I mass of $5.9 \times 10^8 M_{\odot}$. C2 appears to be embedded in the same H I envelope as C1 (see Fig. 3), which has an H I mass of $2.2 \times 10^8 M_{\odot}$. The two clumps are found within an H I tail that is oriented parallel to the gas-poor lenticular galaxy IC 5181. C1 is just reaching the minimum H I mass threshold that is needed to form a long-lived TDG (Bournaud & Duc 2006), meaning that it is likely a transient tidal feature. On the other hand, the H I mass of C2 is well above $10^8 M_{\odot}$. This implies that C2 has the potential to develop into the self-gravitating long-lived TDG. In addition to its H I mass, C2 appears to be at the tip of a tidal tail and is spatially distanced from its assumed parent galaxies. The lack of optical counterparts associated with C2 could indicate that the system is still young with respect to the ongoing interaction event.

The H I cloud C5 is located close to the spiral galaxy NGC 7232B and has an H I mass of $2.2 \times 10^8 M_{\odot}$. Its velocity is $\sim 1876 \text{ km}^{-1}$ similar to the H I gas associated with NGC 7232 and NGC 7233. Its position suggests that it is a result of an interaction between the spiral galaxies NGC 7232B and NGC 7233. The location of C5 and its H I mass (just slightly above $10^8 M_{\odot}$) suggests that this cloud is likely to fall back into the parent galaxies. The H I cloud C6 is located between NGC 7233 and NGC 7232B, and is aligned opposite to C5. C6 has an H I mass of $4.2 \times 10^8 M_{\odot}$ and its peak column density is $\sim 10^{20}$

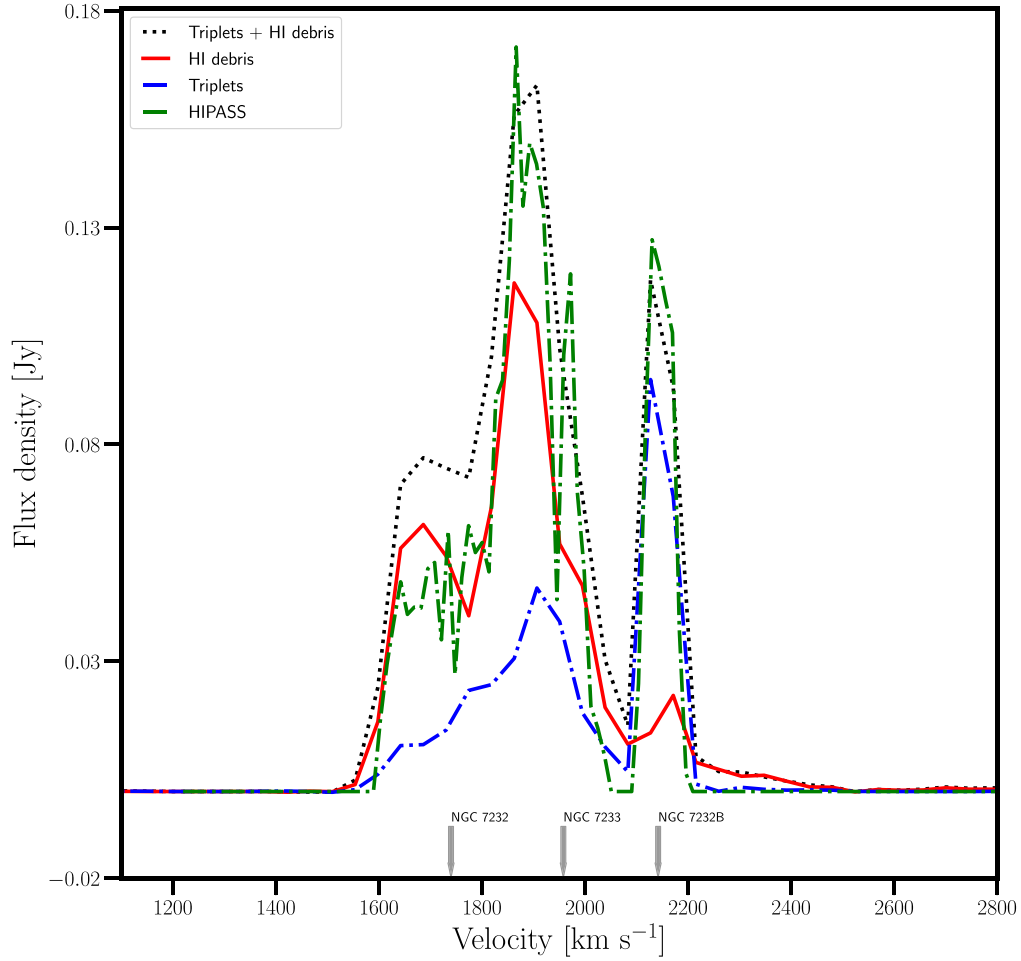


Figure 6. MeerKAT-64 H I global profile of both the NGC 7232/3 triplet and surrounding H I debris (black dotted line), the H I debris alone (red solid line), the galaxy triplet (blue dash-dotted line), and the HIPASS global profile (green dash-dotted line). The HIPASS global profile was extracted after applying a 4σ threshold to the cube.

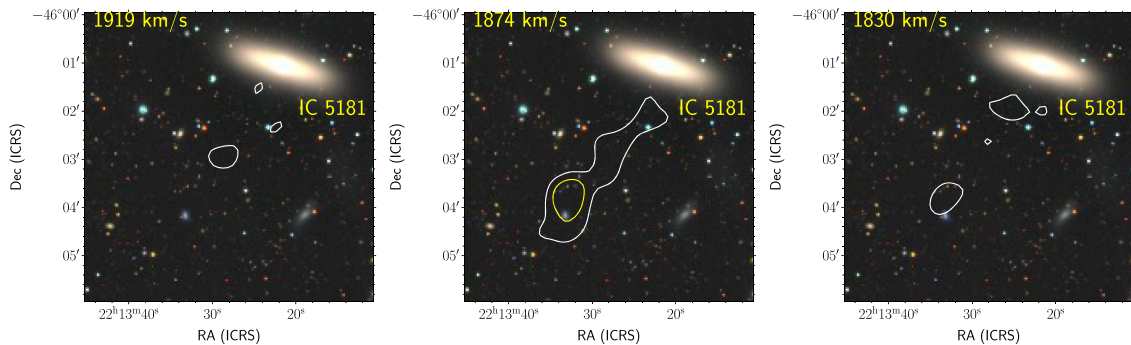


Figure 7. MeerKAT-64 channel maps (primary beam corrected) showing the H I cloud near the gas poor lenticular galaxy IC 5181. The channel maps are overlaid on the DECALS *grz*-band image. The contour colours are in increasing order wit: white ($1 \times 10^{19} \text{ cm}^{-2}$) and yellow ($2.6 \times 10^{19} \text{ cm}^{-2}$), where ($1 \times 10^{19} \text{ cm}^{-2}$) represent a 4σ detection limit in the MeerKAT-64 H I cube. The channel velocity is shown in the top left-hand corner.

cm^{-2} . The proximity of C6 to its parent galaxies implies that the H I cloud is likely to fall back into the parent galaxies (Bournaud & Duc 2006). However, as the interaction between NGC 7232B and NGC 7233 is still actively ongoing, it is also possible that the tail containing C6 might extend further, moving C6 outwards. If this happens and C6 retains or even gains more H I mass, there is a possibility that it

might evolve into a long-lived TDG. As indicated in Bournaud et al. (2007), not all potential TDGs are located at the tip of the tails so the (projected) position of C6 within the tidal tail might be less of a factor in its development. Optical and UV images of show faint blue features within the peak H I column density contours of C6 (see Fig. 8, left-hand panel). These optical features are more pronounced

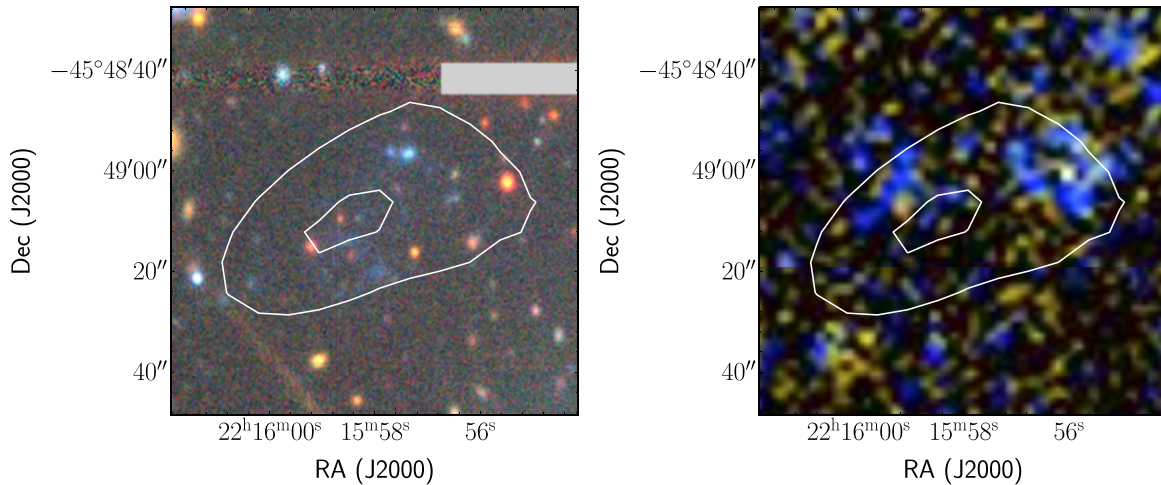


Figure 8. Left-hand panel: MeerKAT-64 integrated H I intensity map of the C6 cloud overlaid on the DECaLS *grz*-band image. Faint blue features within the H I contour could indicate potential *in situ* star formation within the H I cloud C6. Right-hand panel: MeerKAT-64 integrated H I intensity map of the C6 cloud overlaid on a combined *GALEX* NUV/FUV image. The blue colour (FUV) indicates the presence of young stellar populations within the H I contour. The contours are $(5.0, 6.5) \times 10^{20} \text{ cm}^{-2}$.

in UV (see Fig. 8, right-hand panel), thus suggesting the presence of newly formed stars.

5 SUMMARY

We present results from the MeerKAT open time science observations obtained in the *L* band (20 cm) using the 4k-wideband mode (44 km s^{-1} velocity resolution). The aim of the observations was to search for H I emission associated with the NGC 7232/3 galaxy triplet and provide an understanding to the interaction processes between the triplet with its surrounding environment. With a 4σ column density sensitivity of $1 \times 10^{19} \text{ cm}^{-2}$, we find extended H I debris around the NGC 7232/3 galaxy triplet. The newly discovered H I streams extend over ~ 20 arcmin (140 kpc) in projection, and contain an H I mass of $\sim 4.5 \times 10^9 M_{\odot}$, more than 50 per cent of the total H I mass of the galaxy triplet. This result provide a new understanding of where the gas resides within the galaxy triplet and clearly shows how galaxy interactions play an important role in the evolution of galaxies in the NGC 7232/3 galaxy triplet.

The presence of low-column-density H I as well as the distorted morphologies of the triplet galaxies to some extent indicate galaxy interactions. We find that NGC 7232 and NGC 7233 have lost the vast majority of their expected H I mass as opposed to NGC 7232B that seems to have retained its H I. The velocities of each galaxy in the triplet show that NGC 7232B is significantly offset as compared to NGC 7233 and NGC 7232 and that the velocities of the H I debris are consistent with the galaxy pair NGC 7232 and NGC 7233. This suggests that NGC 7232B contains excess H I because it has not really lost a large amount of gas to the intragalactic medium of the triplet, and that most of the extended H I debris gas originated from the galaxy pair. Overall, we find that the galaxy triplet is not H I-deficient. The state of the NGC 7232/3 triplet is analogous to a phase 2 compact group. A similar result has been found for HGC 16 (Jones et al. 2019).

We detected 15 extragalactic H I sources in our data cube and confirmed the detection of six individual H I clouds from previous ASKAP observations. We report one new detection, an H I counterpart to the galaxy LEDA 130765 (WISEA J221403.60–454843.7),

which is possibly a dwarf transition galaxy with an H I mass of $4.8 \pm 1.0 \times 10^7 M_{\odot}$.

Based on the success of our MeerKAT-64 observations, we look forward to a much larger project mapping the star formation and H I gas distribution within nearby galaxy groups to assess their evolutionary state and formation history. The availability of high velocity resolution in future will allow for much advanced kinematic analysis of galaxy dynamics and multiwavelength studies to investigate the interaction processes occurring in the NGC 7232/3 galaxy group.

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DATA AVAILABILITY

The data from this study are available upon request to the corresponding author, Brenda Namumba.

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APPENDIX

A1 H I-detected galaxies in NGC 7232 field

We detect 16 galaxies in H I over a velocity range of ~ 700 – 4000 km s^{-1} . For comparison, Lee-Waddell et al. (2019) studied the H I emission in the NGC 7232 galaxy group over a velocity range of ~ 1500 – 3000 km s^{-1} , while Kleiner et al. (2019) focused on the foreground galaxy IC 5201 and its surroundings (focusing on the velocity range from -20 to 2513 km s^{-1}), both using selected pointings from the first WALLABY Early Science observations. In this section, we briefly describe the general properties of the H I-detected galaxies in our MeerKAT-64 observations. Table A1 summarizes the H I properties of the galaxies and compares with ASKAP (Lee-Waddell et al. 2019) and HIPASS Koribalski et al. (2004). Figs A1–A3 show the integrated maps, velocity field maps, and global profiles of each galaxy:

- (i) The nearby spiral galaxy IC 5201 lies in the foreground of the NGC 7232 group, and was studied in detail by Kleiner et al. (2019) using WALLABY Early Science data. IC 5201 is detected at the eastern edge of our MeerKAT-64 field. We derive an H I mass of $(7.90 \pm 0.04) \times 10^9 M_{\odot}$ using an adopted distance of 13.2 Mpc

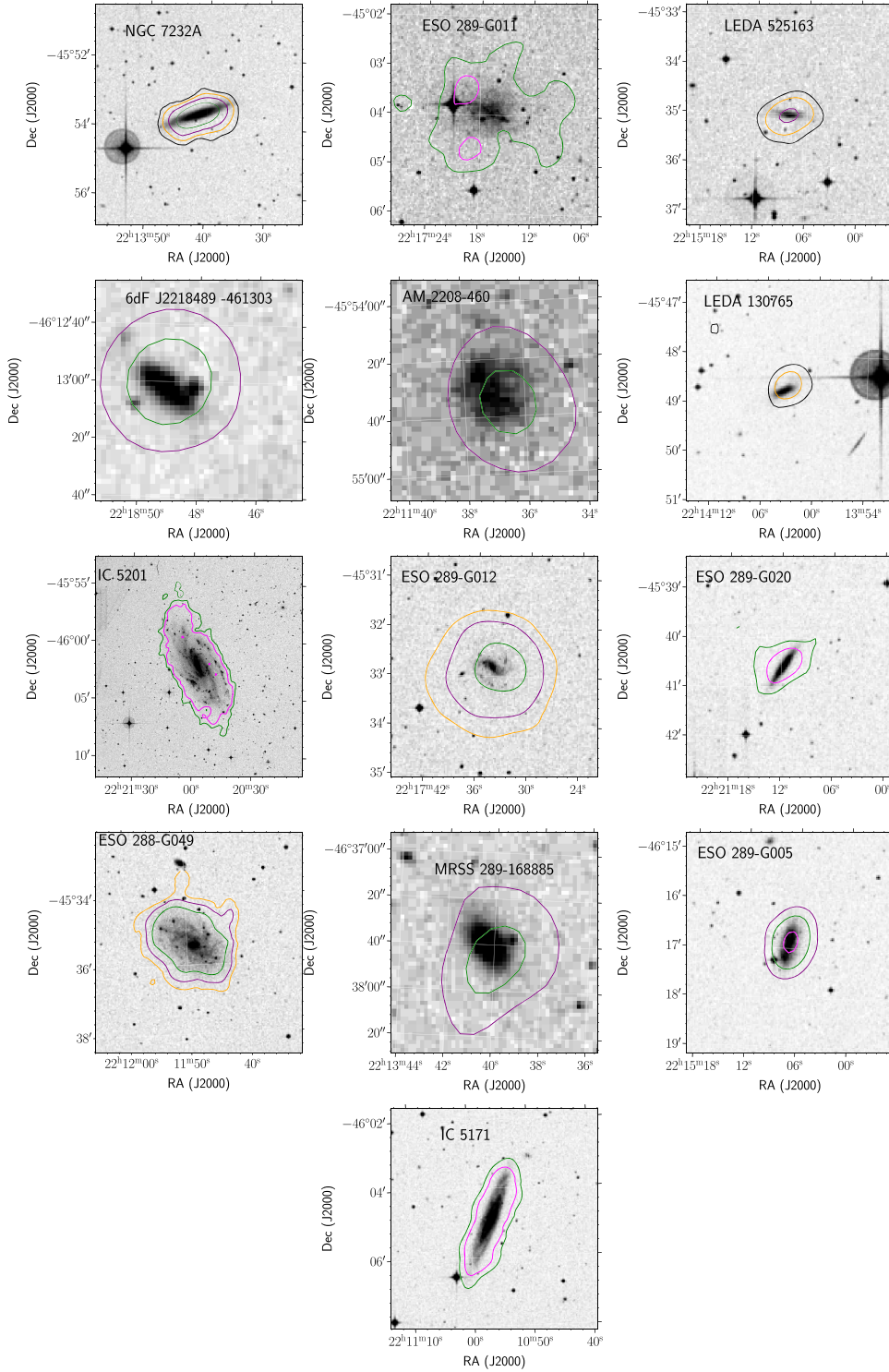


Figure A1. MeerKAT integrated H I column density maps (primary beam corrected) of individual galaxies in the NGC 7232 group overlaid on DSS optical images. The contour levels are $1 \times 10^{19} \text{ cm}^{-2} \times (2, 4, 8, 16, 32)$, where $1 \times 10^{19} \text{ cm}^{-2}$ represent the 4σ detection limit. The colours are in order of increasing column density: black, orange, purple, green, and magenta.

(Kleiner et al. 2019). We measure a systemic velocity of $(914 \pm 52) \text{ km s}^{-1}$.

(ii) ESO 289-G020 is a companion galaxy of IC 5201, also detected by Kleiner et al. (2019). It is an edge-on galaxy with a systemic velocity of $(925 \pm 14) \text{ km s}^{-1}$. We measure an H I mass of

$(1.4 \pm 0.2) \times 10^8 M_{\odot}$ (assuming $D = 13.2 \text{ Mpc}$). Our derived H I mass is in agreement with the ASKAP observations (Kleiner et al. 2019).

(iii) For the spiral galaxy ESO 288-G049, we find a systemic velocity of $(\sim 1950 \pm 15) \text{ km s}^{-1}$ and calculate an H I mass of

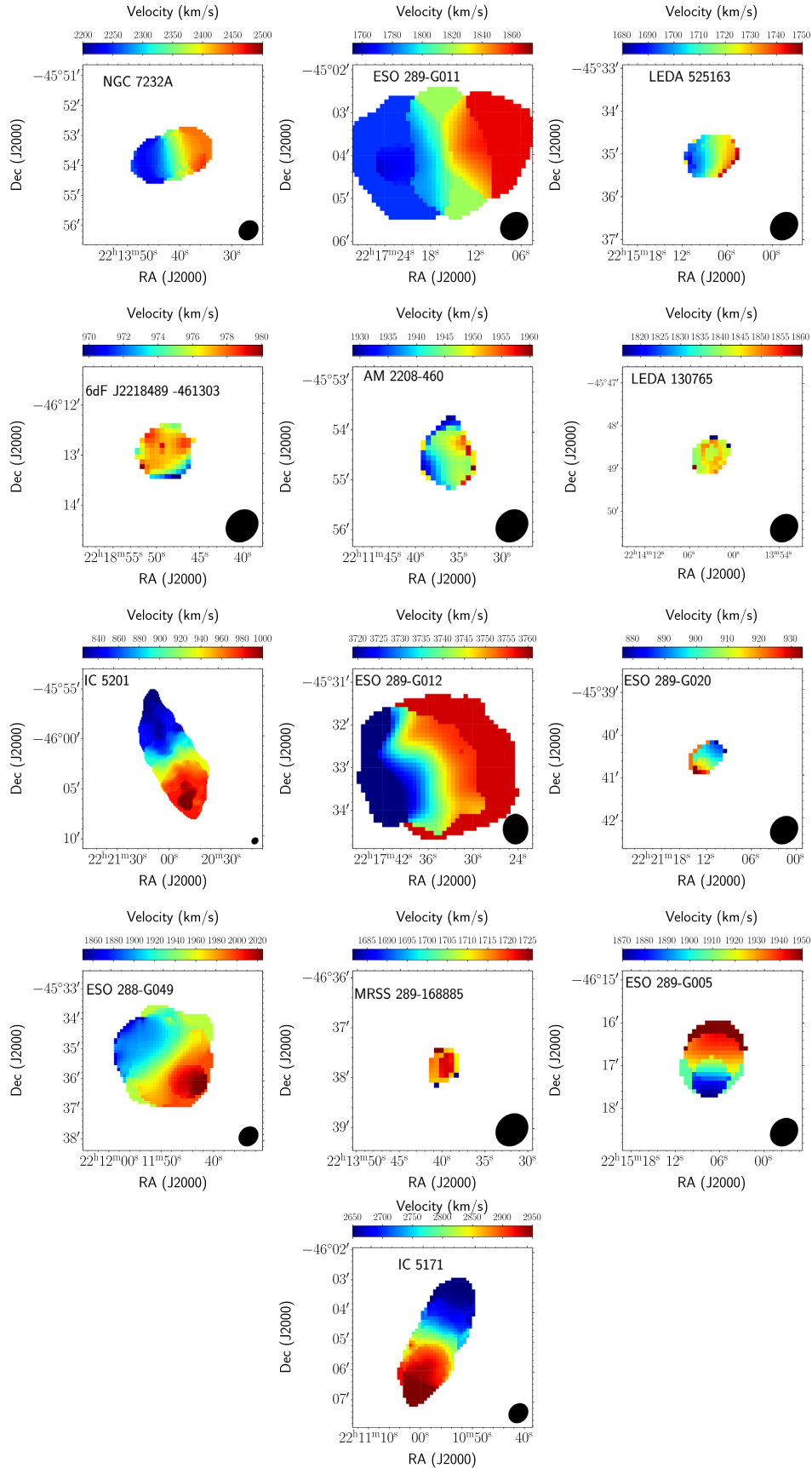


Figure A2. MeerKAT mean H1 velocity fields of individual galaxies in the NGC 7232 galaxy group. The synthesized beam is shown in the bottom right-hand corner.

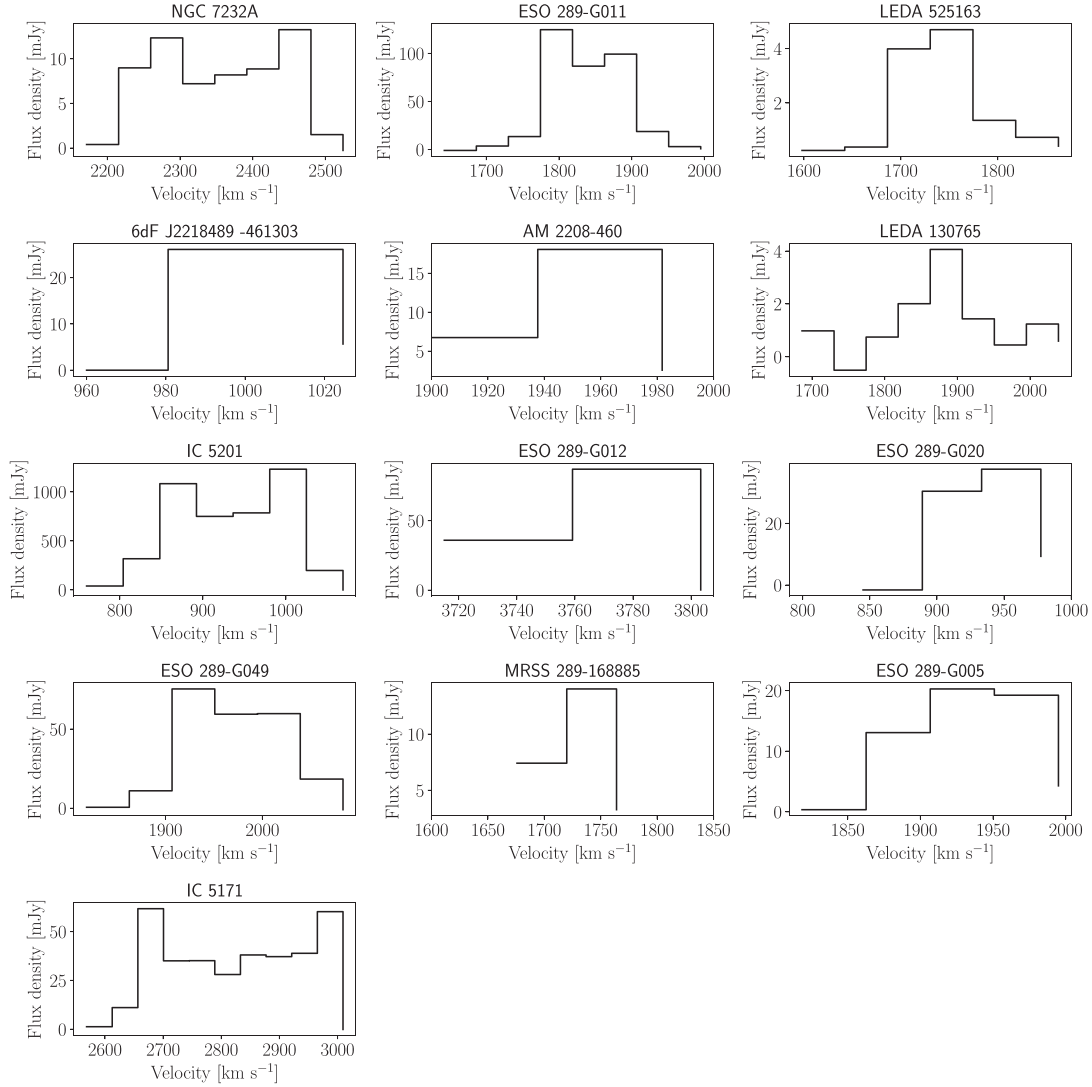


Figure A3. MeerKAT H I intensity profiles of individual galaxies in the NGC 7232 group obtained using SOFIA, i.e. only H I signal deemed reliable is shown.

$(1.3 \pm 0.5) \times 10^8 M_{\odot}$. Lee-Waddell et al. (2019) report an ASKAP H I mass which is ~ 39 per cent less than our MeerKAT value.

(iv) For the edge-on spiral galaxy IC 5171, we find a systemic velocity of $(\sim 2801 \pm 115) \text{ km s}^{-1}$ and an H I mass of $(5.4 \pm 0.3) \times 10^9 M_{\odot}$ ($D = 38.8 \text{ Mpc}$). Following Lee-Waddell et al. (2019), MeerKAT detects ~ 27 per cent more H I mass than ASKAP. The HIPASS H I mass is in good agreement with our MeerKAT results.

(v) For the irregular dwarf galaxy ESO 289–G011, we measure a systemic velocity of $(1816 \pm 36) \text{ km s}^{-1}$. Adopting a distance of 24 Mpc, we derive an H I mass of $(2.10 \pm 0.04) \times 10^9 M_{\odot}$. MeerKAT detects more H I mass when compared with ASKAP (Lee-Waddell et al. 2019).

(vi) The spiral galaxy ESO 289–G005 was first detected in WALLABY Early Science data by Lee-Waddell et al. (2019). We measure a systemic velocity of $(1913 \pm 25) \text{ km s}^{-1}$. An H I mass of $(3.4 \pm 0.4) \times 10^8 M_{\odot}$ is derived ($D = 24 \text{ Mpc}$). Our derived H I mass is in agreement with the ASKAP reported by Lee-Waddell et al. (2019).

(vii) AM 2208–460 is a dwarf galaxy, first detected in H I in WALLABY Early Science data (Lee-Waddell et al. 2019). We measure a systemic velocity of $(\sim 1958 \pm 7) \text{ km s}^{-1}$. No velocity from optical spectroscopy exists in the literature. Adopting a distance of 24 Mpc, we derive an H I mass of $(1.8 \pm 0.3) \times 10^8 M_{\odot}$. Compared to the ASKAP, we derive ~ 50 per cent more H I mass.

(viii) LEDA 525163 is a dwarf irregular galaxy. Here, we measure a systemic velocity of $(1715 \pm 15) \text{ km s}^{-1}$. We derive an H I mass of $(6.1 \pm 1.0) \times 10^7 M_{\odot}$ for a distance of 24 Mpc.

(ix) 6dF J2218489–461303 (PGC 130771) is a foreground dwarf irregular galaxy near IC 5201. We find H I emission centred around $(984 \pm 1) \text{ km s}^{-1}$. We derive an H I mass of $(1.5 \pm 0.2) \times 10^8 M_{\odot}$ ($D = 24 \text{ Mpc}$). Kleiner et al. (2019) reports an H I mass ~ 44 per cent less.

(x) MRSS 289–168885 is possibly a dwarf transition galaxy. Its optical appearance is spheroidal and undisturbed. We find H I emission centred around $(1721 \pm 9) \text{ km s}^{-1}$. Adopting a distance of 24 Mpc, we derive an H I mass of $(8.4 \pm 2.3) \times 10^7 M_{\odot}$.

Table A1. Other H I sources detected with MeerKAT in the NGC 7232 galaxy group.

| Galaxy | Type | α, δ (J2000) (^h m s) | α, δ (J2000) (^d m s) | Optical v_{hel} (km s^{-1}) | HIPASS $v_{\text{H I}}$ (km s^{-1}) | HIPASS $F_{\text{H I}}$ (Jy km s^{-1}) | ASKAP $v_{\text{H I}}$ (km s^{-1}) | ASKAP $F_{\text{H I}}$ (Jy km s^{-1}) | MeerKAT $v_{\text{H I}}$ (km s^{-1}) | MeerKAT $F_{\text{H I}}$ (Jy km s^{-1}) | MeerKAT $M_{\text{H I}}$ ($10^9 M_{\odot}$) |
|---------------------|--------------|-------------------------------------------------|-------------------------------------------------|----------------------------------------------------|---------------------------------------------------|------------------------------------------------------|--------------------------------------------------|-----------------------------------------------------|----------------------------------------------------|-------------------------------------------------------|--------------------------------------------------|
| IC 5201 | SB(rs)cd | 22:20:57.44 | -46:02:09.1 | 914 | - | - | 914 | 177 ± 1 | 914 ± 15 | 194 ± 1.00 | 7.90 ± 0.04 |
| ESO 289-G020 | - | 22:21:11.71 | -45:40:35.5 | 911 | - | - | - | - | 925 ± 14 | 3.60 ± 0.60 | 0.14 ± 0.02 |
| 6dF J2218489-461303 | - | 22:18:48.93 | -46:13:03.1 | 1207 | - | - | 996 ± 2 | 0.6 ± 1 | 984 ± 1 | 1.09 ± 0.21 | 0.15 ± 0.02 |
| ESO 288-G049 | SB(rs)dm | 22:11:48.56 | -45:35:22.7 | 1968 | 1964 | 6.6 | 1964 | 6 ± 1 | 1950 ± 15 | 9.87 ± 0.40 | 0.13 ± 0.05 |
| IC 5171 | SAB(rs)bc | 22:10:56.70 | -46:04:53.3 | 2847 | 2836 | 15.8 | 2831 | 11 ± 2 | 2801 ± 115 | 15.27 ± 0.82 | 5.40 ± 0.30 |
| ESO 289-G011 | IB(s)m | 22:17:16.33 | -45:03:59.4 | 1818 | 1817 | 10.5 | 1821 | 8 ± 2 | 1816 ± 36 | 15.4 ± 0.29 | 2.10 ± 0.04 |
| ESO 289-G005 | - | 22:15:07.16 | -46:16:53.4 | 1885 | - | 11.9 | 1931 | 2.1 ± 0.4 | 1913 ± 25 | 2.52 ± 0.33 | 0.34 ± 0.04 |
| AM 2208-460 | dwarf | 22:11:37.13 | -45:54:31.0 | - | - | - | 1953 | 0.6 ± 0.1 | 1958 ± 7 | 1.19 ± 0.20 | 0.18 ± 0.03 |
| LEDA 525163 | - | 22:15:08.01 | -45:35:01.4 | - | - | - | - | - | 1715 ± 15 | 0.45 ± 0.11 | 0.06 ± 0.01 |
| MRSS 289-168885 | - | 22:13:40.24 | -46:37:39.9 | - | - | - | 1741 | 0.5 ± 0.1 | 1721 ± 9 | 0.62 ± 0.17 | 0.08 ± 0.02 |
| NGC 7232A | SB(rs)ab | 22:13:41.44 | -45:53:37.4 | 2407 | - | - | 2348 | 2.0 ± 0.4 | 2320 ± 78 | 2.56 ± 0.32 | 0.35 ± 0.043 |
| LEDA 130765 | - | 22:14:03.63 | -45:48:43.6 | - | - | - | - | - | 1824 ± 7 | 0.35 ± 0.07 | 0.05 ± 0.01 |
| ESO 289-G012 | dwarf spiral | 22:17:34.29 | -45:32:48.8 | 3792 | - | - | - | - | 3564 ± 15 | 5.46 ± 0.21 | 3.60 ± 0.14 |

Notes. (1) Galaxy name; (2) galaxy type (de Vaucouleurs et al. 1991); (3) RA and Dec. (de Vaucouleurs et al. 1991); and (4) optical v_{hel} (de Vaucouleurs et al. 1991); Columns (5) and (6): Meyer et al. (2004); columns (7) and (8): Lee-Waddell et al. (2019); Columns (9) – (11): this work. For IC 5201 and 6dF J2218489-461303; columns (7) and (8): Kleiner et al. (2019).

(xi) NGC 7232A is an edge-on spiral galaxy (type SA), not catalogued in HIPASS. We find H I emission centred around $(2320 \pm 78) \text{ km s}^{-1}$. We derive an H I mass of $(3.5 \pm 0.4) \times 10^8 M_{\odot}$ ($D = 24$ Mpc), similar to the ASKAP (Lee-Waddell et al. 2019). Optical images of NGC 7232A show it to be symmetric and undisturbed.

(xii) LEDA 130765 (WISEA J221403.60-454843.7), not mentioned by Lee-Waddell et al. (2019), is possibly a dwarf transition galaxy located just north of the extended tidal debris seen in Fig. 5. Its optical appearance is lenticular and undisturbed. We find H I emission centred around $(1824 \pm 7) \text{ km s}^{-1}$ and measure an H I mass of $(4.8 \pm 1.0) \times 10^7 M_{\odot}$ ($D = 24$ Mpc).

(xiii) The galaxies LEDA 519373 and 519377 (around 2008 km s^{-1}) are located just south of the lenticular galaxy IC 5181 near the debris at the south-western end of the extended H I filament. We measure a total H I flux of $F_{\text{H I}} \sim 1.83 \pm 0.05 \text{ Jy km s}^{-1}$ in that region.

(xiv) ESO 289-G012 is a spiral galaxy behind the NGC 7232 group. We measure a systemic velocity of $(3564 \pm 15) \text{ km s}^{-1}$ and calculate an H I mass of $(3.6 \pm 0.1) \times 10^9 M_{\odot}$ ($D = 53$ Mpc).

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