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# The Murchison Widefield Array Transients Survey (MWATS). A search for low-frequency variability in a bright Southern hemisphere sample

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## ABSTRACT

We report on a search for low-frequency radio variability in 944 bright ( $>4$  Jy at 154 MHz) unresolved, extragalactic radio sources monitored monthly for several years with the Murchison Widefield Array. In the majority of sources, we find very low levels of variability with typical modulation indices  $<5$  per cent. We detect 15 candidate low-frequency variables that show significant long-term variability ( $>2.8$  yr) with time-averaged modulation indices  $\bar{M} = 3.1\text{--}7.1$  per cent. With 7/15 of these variable sources having peaked spectral energy distributions, and only 5.7 per cent of the overall sample having peaked spectra, we find an increase in the prevalence of variability in this spectral class. We conclude that the variability seen in this survey is most probably a consequence of refractive interstellar scintillation and that these objects must have the majority of their flux density contained within angular diameters less than 50 milliarcsec (which we support with multiwavelength data). At 154 MHz, we demonstrate that interstellar scintillation time-scales become long ( $\sim$ decades) and have low modulation indices, while synchrotron-driven variability can only produce dynamic changes on time-scales of hundreds of years, with flux density changes less than one milli-jansky (without relativistic boosting). From this work, we infer that the low-frequency extragalactic southern sky, as seen by SKA-Low, will be non-variable on time-scales shorter than 1 yr.

**Key words:** radio continuum: galaxies – radio continuum: ISM – radio continuum: transients.

## 1 INTRODUCTION

The phenomenon of low-frequency variability (LFV) was first clearly confirmed by Hunstead (1972). In that study, four bright extragalactic variable radio sources were identified that showed amplitude modulation of up to 20 per cent on time-scales under 5 yr at 408 MHz. At that time the variations were difficult to reconcile as they broke the Compton catastrophe limit (Kellermann & Pauliny-Toth 1969), meaning that the mechanism for variability

could not be intrinsic to the source. Further studies at 408 MHz confirmed that the phenomenon was common, with 25 per cent of compact sources and 51 per cent of flat spectrum sources displaying dynamic properties (McAdam 1979; Fanti et al. 1981; Spangler et al. 1989; see also Jauncey et al. 2016, for a review).

New low-frequency interferometers such as the Murchison Widefield Array (MWA; Tingay et al. 2013) and the Low Frequency Array (LOFAR; van Haarlem et al. 2013) allow for greatly improved spatial, temporal, and spectral coverage. Hence, we can now efficiently probe large samples of low-frequency sources for long-time-scale dynamic changes.

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Scaling up the Hunstead (1972) results to the abundance of radio sources now routinely being detected by these new instruments (hundreds of thousands), one might assume the extragalactic low-frequency sky to be highly variable. In this paper, we conduct a comprehensive analysis of the abundance of low-frequency variables as seen by the MWA and on the basis of this make predictions for the Square Kilometre Array (Dewdney et al. 2009).

The mechanism driving the variability in the sources identified by Hunstead (1972) has been determined to be interstellar scintillation (ISS) caused by the turbulent and ionized interstellar medium (ISM; see Rickett et al. 1984). Scintillation can be divided into several different subclasses that are a function of source angular size, observing frequency, and line of sight through the ISM which will influence the amplitude and time-scale of variability (see Lazio et al. 2004, for a review). Specifically, refractive scintillation is bending of different ray paths through the clumpy ISM, while diffractive scintillation is interference between different ray paths.

LFV is now attributed to refractive interstellar scintillation (RISS) and is prominent when the source sizes are  $\sim 1$  milliarcsec across and the observing frequency is low  $< 500$  MHz (Rickett 1986). At higher frequencies intraday variability (IDV; Dennett-Thorpe & de Bruyn 2002; Bignall et al. 2003) can be thought of as the higher frequency ‘sibling’ propagation effect to the LFV. The IDV can cause large intensity fluctuations on short time-scales ( $< 24$  h) and is most prominent at frequencies  $\sim 5$  GHz. Sweeping from high (GHz) to low (MHz) frequencies, we transition from the weak scattering regime (less than 1 rad of phase variation, with a modulation index  $< 1$ ) to the strong ( $\gg 1$  rad of phase variation). In the strong regime if the source is compact enough compared to an isoplanatic patch (typically only for pulsars) it will exhibit diffractive scintillation, but larger sources such as active galaxies exhibit refractive scintillation with slow (time-scale  $\propto \nu^{-2.2}$ ) lower amplitude (modulation index  $\propto \nu^{0.57}$ ) fluctuations (Lazio et al. 2004).

Interplanetary scintillation (IPS; Hewish, Scott & Wills 1964) requires a contrast in density of the intervening medium, this time in the interplanetary medium, which is driven by inhomogeneities in the solar wind. Although IPS is typically observed at small solar elongations, it has been detected serendipitously by the MWA in night-time observing (Kaplan et al. 2015) with flux density variations measured on the time-scale of seconds. This has been further exploited by Morgan et al. (2018) and Chhetri et al. (2018) to catalogue hundreds of sources with a subarcsec structure across the sky. A major finding of this work is that almost all peaked-spectrum sources show strong IPS, meaning that they are compact on subarcsec scales.

A number of surveys have now studied the time-domain radio sky specifically searching for (and studying) propagation effects at a range of frequencies. For example, at GHz frequencies, the Micro-Arcsecond Scintillation-Induced Variability survey (MA-SIV; Lovell et al. 2003; Lovell et al. 2008; Pursimo et al. 2013) found that over 50 per cent of flat spectrum sources show significant variability at 5 GHz. The authors do report, however, that a higher abundance of these variables is detected in their weak sample ( $\sim 0.1$  Jy) as opposed to their strong sample ( $\sim 1$  Jy). The faster, more intense variability seen at 5 GHz (as discussed above) is driven by weak ISS (see Jauncey et al. 2016, for a review).

At frequencies comparable to this work (154 MHz), some observational studies have reported that the prevalence of LFV is more rare than at 408 MHz. For example, McGilchrist & Riley (1990) show that only 1.1 per cent of a sample of 811 sources with flux densities  $> 0.3$  Jy show fractional variability  $> 15$  per cent on time-scales of 1 yr at 151 MHz (also see Riley 1993). In contrast, Slee &

Siegmán (1988) report that 13.3 per cent of a sample of 412 sources  $> 1$  Jy show fractional variability  $> 15$  per cent at 160 MHz. The total length of the Slee & Siegmán (1988) study was, however, longer (up to 14 yr), although variability was still reported on month time-scales. In both cases, however, the abundance of variables is lower than reported at 408 MHz ( $\sim 25$  per cent; see McAdam 1979 and Fanti et al. 1981). In more recent work, Rowlinson et al. (2016) probed an unbiased sample of predominantly extragalactic radio sources with flux densities  $\geq 0.5$  Jy over a period of 1 yr with the MWA at 182 MHz. In this study, over 100 h of observations were analysed to produce more than ten thousand wide-field images ( $600 \text{ deg}^2$  each) containing around 5000 sources. For this sample, low levels of variability were reported with modulation indices typically below 10 per cent (Rowlinson et al. 2016). In comparison, Bell et al. (2014) used the pathfinder MWA instrument to study the LFV and found 1.5 per cent of sources with flux densities  $\geq 3$  Jy showed significantly fractional variability with variations  $> 10$  per cent on time-scales of 1 yr.

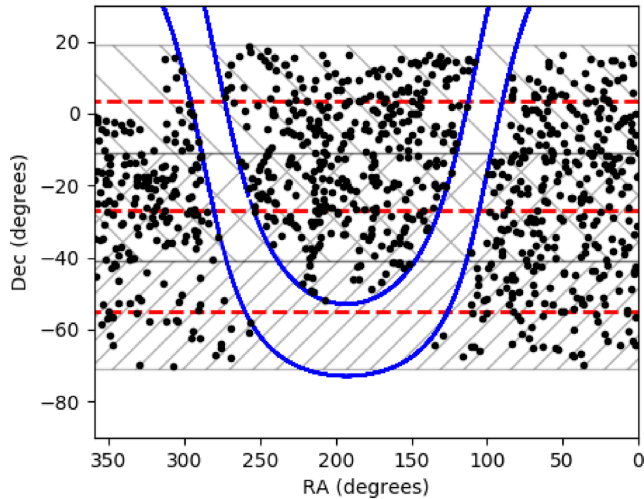
Clearly, the abundance of variable sources is lower at low frequencies, which is supported by the theoretical frequency dependence of RISS. But with the recent renaissance in low-frequency radio astronomy the time was right to revisit the field of LFV and pursue a systematic census of its abundance prior to SKA-Low operations. The MWA also has exceptional spectral coverage at low frequencies, so for the first time we could explore the detailed spectral energy distributions (SEDs) of variables. We focused this survey on a pre-selected sample of 944 compact extragalactic sources distributed over the Southern hemisphere and probed time-scales up to 2.8 yr.

This paper uses data taken for the Murchison Widefield Array Transients Survey (MWATS; P.I. M. Bell). MWATS was an observing campaign aimed at sampling the dynamic radio sky on time-scales of minutes to years over the Southern hemisphere at 154 MHz. MWATS has already provided time-domain measurements of objects such as pulsars (Bell et al. 2016), exoplanets (Murphy et al. 2015), the ionosphere (Loi et al. 2015a; Loi et al. 2015b; Loi et al. 2016), and low-mass stars. The end goal of MWATS was to provide a complete time-domain Southern hemisphere census of known and previously unknown objects (i.e. transients) prior to SKA-Low operations.

In this paper, we revisit the topic of LFV through the lens of our survey(s). In Section 2, we describe the survey specifications and sample selection. In Section 3, we describe the calibration, imaging, and light-curve extraction procedure. In Section 4, we examine the time-domain, spectral, and multiwavelength properties of these objects to identify possible variables. In Section 5, we discuss the physical interpretation of our results and compare them with past and upcoming surveys. Unless otherwise quoted, all uncertainties correspond to  $1\sigma$  confidence. We define the relationship between flux density  $S_\nu$ , frequency  $\nu$ , and spectral index  $\alpha$  as  $S_\nu \propto \nu^\alpha$ .

## 2 SURVEY SPECIFICATIONS

The observations for this survey commenced on 2013 September 16 and ended on 2016 June 25, covering a time span of 2.8 yr. In total there were 24 nights of data, each separated by more than 10 d, giving a total of 5639 images. Each image covered an area of sky approximately  $1000 \text{ deg}^2$ . Observations were centred at 154 MHz with an observing bandwidth of 30.72 MHz. The bandwidth was divided into 768 channels of width 40 kHz. The correlator integration cycle was either 0.5 or 2 s; it was increased in later observations (after 2014) to reduce disc usage.



**Figure 1.** The spatial distribution of objects in our sample. The hatched areas show the drift scan regions that are centred on declinations of  $+1.6^\circ$ ,  $-26^\circ$ , and  $-55^\circ$  (red dashed lines). Sources in the Galactic plane are missing as they are not included in the GLEAM EGC (Hurley-Walker et al. 2017). This exclusion zone is defined by the blue lines where  $b = \pm 10^\circ$ .

**Table 1.** Properties of observations.

Property	Value
Total images	5639
Centre frequency	154 MHz
Bandwidth	30.72 MHz
Channelization	40 kHz
Integration time per image	112 s
Correlation time	0.5 s or 2 s
Image size (pixels)	$4000 \times 4000$
Pixel size	$0.75 \text{ arcmin} \times 0.75 \text{ arcmin}$
Beam diameter (FWHM)	2.25 arcmin
Briggs weighting	-1
UV selection ( $k\lambda$ )	$>0.03 k\lambda$
Declinations	$+1.6^\circ, -26^\circ, -55^\circ$
Cadence	minutes, months, and years
Typical image noise (extragalactic pointing)	-
	$20 \text{ mJy beam}^{-1}$
Typical image noise (Galactic pointing)	-
	$100 \text{ mJy beam}^{-1}$
Survey duration	2.8 yr

The survey strategy utilized drift scanning to repeatedly observe three meridian points at declination centres of  $-26^\circ$  (zenith),  $+1.6^\circ$ , and  $-55^\circ$  (see Fig. 1). Owing to the large field of view of the MWA this strategy gave uniform coverage at 154 MHz between declinations of  $-70^\circ$  and  $+20^\circ$ . Each observation consisted of a 112 s integration at each of the meridian points in turn. The mean time difference between all observations regardless of declination was 2 min. For the same declination strip it was 6 min, due to declination switching. Observations were always obtained at night and seasonal sky rotation allowed us to cover the complete RA range per year. This had an impact on the cadence of observations for the objects in our sample. A given source was typically observable for 3–4 months per year depending on its proximity to the drift scan centres. A summary of the specifications is given in Table 1.

## 2.1 Sample selection

Using the GaLactic and Extragalactic All-sky Murchison Widefield Array (GLEAM) survey (Wayth et al. 2015), we extracted a biased sample of bright ( $>4 \text{ Jy}$  at 151 MHz) and possibly compact radio sources defined via visual and spectral inspection from the Extragalactic Catalogue (EGC; Hurley-Walker et al. 2017, White et al., in preparation). We take all components in the GLEAM EGC that are brighter than 4 Jy (total of 1879 objects). From visual inspection, we removed objects that had the following properties:

- (i) were resolved into two or more components in images from the NRAO VLA Sky Survey (NVSS; Condon et al. 1998) or the Sydney University Molonglo Sky Survey (SUMSS; Mauch et al. 2003);
- (ii) were confused in the MWA beam, i.e. there were two unresolved NVSS or SUMSS sources within the error ellipse of the MWA beam.

This left a sample of 1143 objects.

Within our survey area, we then constructed light curves for these objects, which gave a final sample of 944 objects. This lower sample size is due to the different footprints of MWATS and the GLEAM surveys. GLEAM covers a declination range  $-90^\circ < \delta < +30^\circ$  whereas MWATS covers  $-70^\circ < \delta < +20^\circ$ , leading to a smaller overlap region and hence smaller sample size. The spatial distribution of our 944 objects is shown in Fig. 1. Note, the GLEAM EGC does not report sources that fell within a number of exclusion zones (see Hurley-Walker et al. 2017, for full details of exclusion regions). For example, sources within  $10^\circ$  of the Galactic plane, or within  $5.5^\circ$  and  $2.5^\circ$  of the Large or Small Magellanic Clouds, respectively, were excluded. As our sample was derived from the GLEAM EGC, we consequently do not consider sources in these regions.

## 3 DATA REDUCTION

### 3.1 Calibration

For each night of observations a single 112 s observation centred on a bright, well-modelled (not necessarily unresolved) source was obtained for phase calibration purposes. The models used for calibration were extracted from previous radio surveys of the Southern hemisphere, namely SUMSS or the VLA Low-frequency Sky Survey (VLSS; Cohen et al. 2007). They are drawn from the same set of models used for calibration of the GLEAM survey. The model images were inverse Fourier transformed to generate a model set of visibilities. A single time-independent, frequency-dependent amplitude and phase calibration solution was derived from this model and compared with the calibrator observation visibilities. These gain solutions were then applied to the nightly drift scan observations prior to imaging.

### 3.2 Imaging

The imaging procedure was almost identical to that done in Bell et al. (2016) and very close to the processing performed for the GLEAM survey. The major change for this work (compared with Bell et al. 2016) was the addition of the peeling and deep CLEAN steps described below. For each of the observations, we did the following steps:

- (i) *Flagging*: The visibilities were flagged for radio frequency interference using the AOFLAGGER algorithm (Offringa, van de



Gronde & Roerdink 2012) and converted into CASA measurement set format using the MWA preprocessing pipeline COTTER. Approximately 1 per cent of the visibilities were removed at this stage (see Offringa et al. 2015, for more details).

(ii) *Bandpass calibration*: Phase solutions derived from the calibrator observation were applied to the visibilities (as discussed above).

(iii) *Peeling*: Bright sources ( $>50$  Jy) that fell within the sidelobe regions were peeled and removed to aid in both self-calibration (described below) and image fidelity. Peeling involves using a model of the bright source and a self-calibration cycle to adequately subtract that source from the visibilities (see Hurley-Walker et al. 2017, for further details).

(iv) *Shallow CLEAN*: The visibilities were deconvolved and CLEANed with 2000 iterations using the WSCLEAN algorithm (Offringa et al. 2014). The full 30.72 MHz bandwidth was imaged using the Multi-Frequency Synthesis (MFS) WSCLEAN algorithm. An RMS noise measurement  $\sigma_N$  was taken from the images to ascertain an appropriate CLEAN threshold for post-self-calibration imaging and deconvolution.

(v) *Self-calibration*: The CLEAN component model was inverse Fourier transformed for self-calibration purposes. A new set of phase and amplitude calibration solutions was derived from this model and applied to the data.

(vi) *Final flagging*: A further round of flagging was performed to tackle any residual interference or erroneous visibilities. The same flagging procedure was applied as discussed above. Approximately 0.1 per cent of the remaining visibilities were flagged at this stage.

(vii) *Deep CLEAN*: The visibilities were then deconvolved and CLEANed. An image size of  $4000 \times 4000$  with a pixel width 0.75 arcmin and robust parameter of  $-1$  (close to uniform weighting) was used (see Table 1 for further details). For images centred away from Galactic plane ( $|b| > 15^\circ$ ), we CLEANed with an unrestricted number of CLEAN components down to a cut-off of  $3 \times \sigma_N$ . For images centred closer to the Galactic plane ( $|b| < 15^\circ$ ), we CLEANed more deeply with a cut-off of  $1 \times \sigma_N$ . This was because there is a larger amount of diffuse material at low Galactic latitudes and more CLEAN components were typically required to reduce the noise levels.

(viii) *Primary beam correction*: A primary beam correction was applied to create Stokes I images; see Offringa et al. (2014) and Sutunjo et al. (2015) for further details. We also produced images in U, Q, V, XX, and YY, although only the Stokes I images were used for the analysis in this paper.

### 3.3 Light-curve extraction

All images for this survey were processed through the Variable and Slow Transients pipeline (VAST; Murphy et al. 2013). The VAST pipeline performs the following processing on each of the images:

(i) *Background estimation*: We use the BANE (Hancock, Trott & Hurley-Walker 2018) algorithm to provide accurate background estimation across each image.

(ii) *Source extraction*: We use the AEGEAN (version 1.9.5; Hancock et al. 2012) source finding algorithm to identify all sources above a local significance of  $5\sigma$ . We restricted source finding to within  $15^\circ$  of the meridian pointing centres. This was to minimize possible variability in flux density arising from uncertainty in the primary beam, which is more pronounced further away from the pointing centres (see Loi et al. 2015a, for discussion).

(iii) *Flux density scale correction*: We use the GLEAM EGC to correct the flux density scale of all of our images. For a given image,

all extracted sources that are within  $20^\circ$  of the pointing centre and brighter than 40 mJy are cross-matched with the GLEAM EGC. For that image, we then calculated a linear fit to the GLEAM versus MWATS flux density points in log space (with zero intercept). The gradient of this log-fit provides a single multiplicative factor, which is applied to each of the images to bring them on to the required flux density scale.

(iv) *Positional correction*: We apply a position-dependent correction for each of the flux density measurements. As the ionosphere can produce positional offsets as a function of time (see Loi et al. 2015a), we use the GLEAM EGC as a reference frame for corrections. The position corrections for an image were determined by matching the sources extracted from the image with the GLEAM EGC. The matching radius was  $0.1^\circ$ . The positional offsets were modelled using a radial basis function and the model offsets were applied to all sources in the image. This is the same process as the image warping described in Hurley-Walker et al. (2017), but it is applied to the extracted sources directly.

(v) *Light-curve generation*: For a source at position RA and Dec., we collate all flux density measurements from all images that are within a diameter of 0.6 arcmin (in RA and Dec.) to form light curves. The flux density measurements were obtained from images averaged over the full 30.72 MHz bandwidth.

## 3.4 Variability analysis

### 3.4.1 Raw light curves

Using all source flux density measurements regardless of time-scale, we calculated the modulation index  $M$ , which is defined as

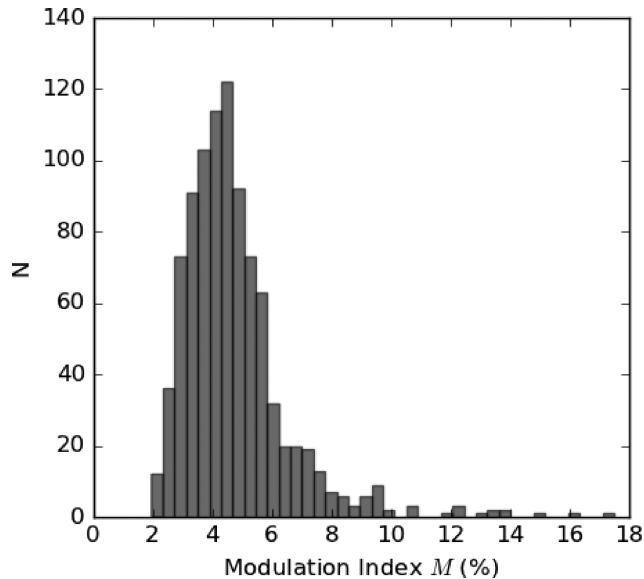
$$M = \sigma / \bar{S}, \quad (1)$$

where  $\sigma$  is the standard deviation and  $\bar{S}$  is the arithmetic mean of the flux density measurements. We refer to these data as the raw light curves.

### 3.4.2 Time-averaged light curves

As we are principally concerned with long-time-scale variations, for each source we combined all measurements on time-scales less than 10 d and recalculated a time-averaged modulation index  $\bar{M}$ . The value of 10 d was chosen for our averaging time-scale as it was just below the minimum time separation between observing nights. Hence, this gave per-observing-night flux density measurements over our observing campaign.

We also calculated a linear least-squares fit to the time-averaged data. From the linear fit, we used the gradient  $\nabla_F$  and fitting error  $\sigma_F$  to calculate a significance of the gradient  $\nabla_S = \nabla_F / \sigma_F$ . We identified candidate variable objects as having  $|\nabla_S| > 3$  and  $\bar{M} > 3$  per cent. We assessed the appropriateness of a simple linear fit by calculating the  $\chi^2$  goodness-of-fit parameter using (a) a flat model (mean of the flux density measurements); (b) the linear fit (discussed above); and (c) a second-order polynomial fit. Using these values, we calculated a  $p$ -value using an F-test (Bevington 1969) comparing the linear fit to the flat model, which we refer to as P-slope. We also compared the polynomial fit to the linear fit, which we refer to as P-quad. A large value of the P-slope (e.g. 0.99) would show that the linear fit to the data is the preferred model when compared with the flat model. A large value of the P-quad would show a preference for a second-order polynomial fit.



**Figure 2.** A histogram of the modulation indices of all raw light curves. We find a median modulation index of  $M = 4.4 \pm 1.8$  per cent.

## 4 RESULTS

### 4.1 Raw light curves

Fig. 2 shows the modulation index  $M$  for all raw light curves. As these light curves had not been averaged in time they were more susceptible to short-time-scale instrumental, measurement, or calibration error. Visual inspection of the small number of high- $M$  light curves ( $M > 10$  per cent) was characterized by the presence of erroneous measurements. Investigation of these points revealed that they were usually associated with a poor-quality image or an incorrect flux density correction, resulting in an artificially large modulation index. We do not therefore believe the tail of this distribution to be caused by physically real variability. Despite these outlier cases, we still found remarkably low levels of variability in the majority of sources (on short time-scales) with a median modulation index of  $M = 4.4 \pm 1.8$  per cent.

### 4.2 Time-averaged light curves

Fig. 3 shows a histogram (right-hand panel) of the modulation indices formed from the time-averaged light curves. We find a median  $\bar{M}$  of  $2.1 \pm 2.0$  per cent, which again is very low. The top panel shows a histogram of the significance of the gradients  $\nabla_S$  of the light curves. In general, the light curves are exceptionally flat with a mean gradient of  $\langle \nabla_F \rangle = -0.22 \pm 3.2 \times 10^{-4}$  Jy per day.

Examining the high- $\bar{M}$  light curves again revealed a small number of erroneous measurements associated with a particular epoch (or night), which had driven up the average modulation index. We could have algorithmically filtered these points out, but with the small number of cases it was easy to visually inspect them and rule them out as being instrumental.

The interesting part of the parameter space is the region where sources have a significant gradient  $|\nabla_S| > 3$  and also an average modulation index  $\bar{M} > 3$  per cent (see Fig. 3). Within this region, we find a set of 15 plausible candidate variables. These are discussed in Section 4.3 with the light curves and SEDs shown in Figs 4–8. On the upper panel of the light curves, we show the raw flux density

measurements plotted sequentially. The lower panel shows the flux density as a function of calendar date, with the linear fit (red line). These results are also summarized in Table 2.

For the majority of these sources (13/15) P-slope  $> 0.90$ , meaning that a linear fit is preferred. Two of the sources (GLEAM J041022–523247 and GLEAM J102820+151129) have P-quad  $> 0.95$  (and both greater than P-slope) meaning that a quadratic fit is more appropriate.

The SEDs (right of the light curves) show the 20 point GLEAM EGC spectrum with red circles together with their corresponding errors. An orange line denotes a single power-law fit to all data points. If the objects are identified as having spectral curvature in Callingham et al. (2017) then a black line represents the best fit to the data using the fitting method described in Callingham et al. (2015). For GLEAM J224704–365746 and GLEAM J010838+013454, there is obvious spectral flattening above the GLEAM band that corresponds to the emission from the source becoming core-dominated. We fit these sources with a broken power law, with the higher frequency power law displayed as a dashed line. Archival measurements are represented as follows:

- (i) Purple upward-pointing triangle: 74 MHz, VLA Low-frequency Sky Survey redux (VLSSr; Lane et al. 2014);
- (ii) Blue square: 150 MHz TIFR GMRT Sky Survey (TGSS-ADR1; Intema et al. 2017);
- (iii) Green leftward-pointing triangle: 408 MHz, Molonglo Reference Catalogue (MRC; Large et al. 1981);
- (iv) Brown rightward-pointing triangle: 843 MHz, SUMSS;
- (v) Navy downward-pointing triangle: 1.4 GHz, NVSS;
- (vi) Green diamond: 4.8, 8.8, and 20 GHz, Australia Telescope 20 GHz survey (AT20G; Murphy et al. 2010).

We also made use of Very Long Baseline Interferometry (VLBI) measurements, if available. This was to ascertain the compact nature of the objects. In particular, we made use of observations from the VLBI Space Observatory Programme (VSOP; Hirabayashi et al. 1998), the International Celestial Reference System (ICRF; Arias et al. 1995), and the Very Long Baseline Array (VLBA; Napier, Thompson & Ekers 1983).

### 4.3 Notes on specific variable sources

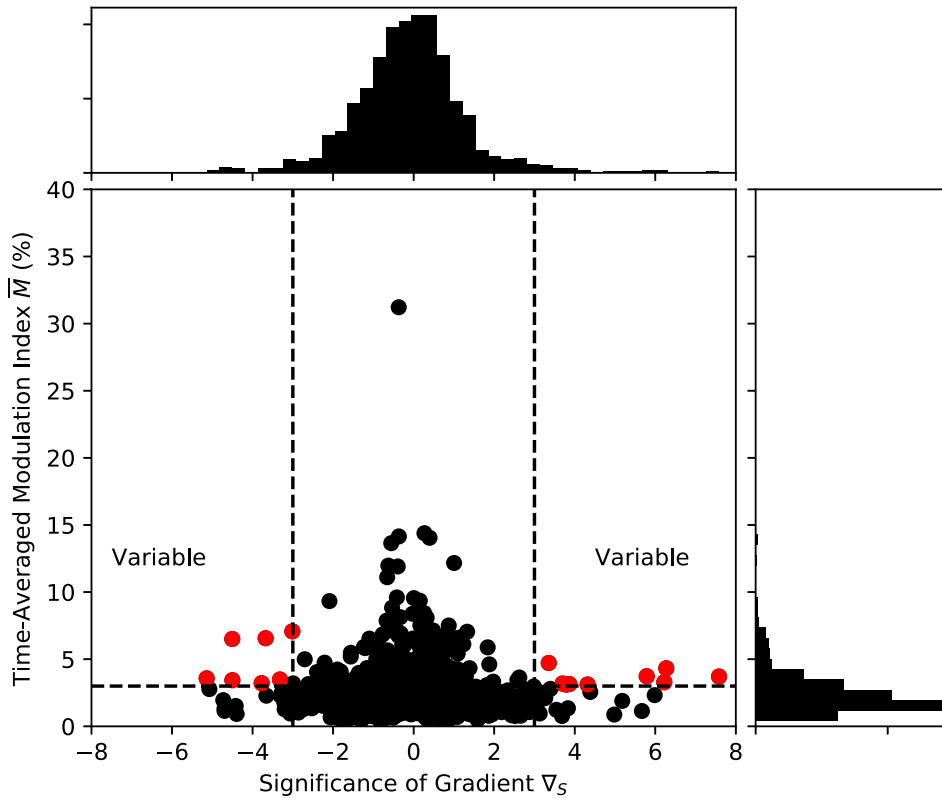
We now discuss the individual properties of the sources that showed significant low-frequency radio variability over the period covered by our study. For a visual comparison, in Fig. 9 we show an example non-variable source (GLEAM J065813 – 654451) that has an averaged modulation index  $\bar{M} = 0.9$  per cent and  $\nabla_S = -0.1$ .

#### 4.3.1 GLEAM J010838+013454 (PKS B0106+013)

This is a well-studied bright Quasar/Blazar (J0108 + 0135; Masaro et al. 2009) with redshift  $z = 2.107$  (Hewett, Foltz & Chaffee 1995) and emission from radio to gamma-rays (Healey et al. 2008). The variability detected through this survey is slow and moderate ( $\bar{M} = 3.1$  per cent). The SED is characterized by a steep power law at frequencies below 500 MHz. Examination of data in the NASA/IPAC Extragalactic Database<sup>1</sup> (NED) shows a flat SED between 1 and 100 GHz, and strong evidence for variability.

This object is slightly too far north to have been observed by the AT20G survey, but it is included in the ICRF and VSOP surveys.

<sup>1</sup><https://ned.ipac.caltech.edu>



**Figure 3.** A plot of averaged modulation index  $\bar{M}$  versus significance of gradient ( $\nabla_S$ ) for all sources in our sample. We find very low levels of long-term variability with a median  $\bar{M} = 2.1 \pm 2.0$  per cent. Most sources remain statistically flat over our measurement time-scale (i.e.  $|\nabla_S| < 3$ ). Red circles denote the candidate variable sources that have a significant gradient  $|\nabla_S| > 3$  and a time-averaged modulation index  $\bar{M} > 3$  per cent.

Using VSOP observations Dodson et al. (2008) give a core size of 0.2 mas at 5 GHz. It is detected as a variable in the MASIV survey with a modulation index of 1.9 per cent on time-scales of 4 d at 4.8 GHz (Koay et al. 2011). In a 2.5 yr monitoring campaign with the Australia Telescope Compact Array, Tingay et al. (2003) report modulation indices of 23, 11, 4, and 4 per cent at 8.6, 4.8, 2.5, and 1.4 GHz, respectively.

#### 4.3.2 GLEAM J032320+053413 (PKS B0320+055)

Best et al. (2003) identify this object as a radio galaxy at redshift  $z = 0.1785$ . They also note that it is a compact ( $< 0.2$  arcsec) radio source, with a high-ionization spectrum. A 1.6 GHz VLBI image made by Liu et al. (2007) in 2006 March shows a strong extended component with a weaker extended component separated by 123 mas. Liu et al. (2007) note that both of these components are probably lobe emission and that their lobe separation is consistent with a Compact Symmetric Object (CSO) radio source. The light curve shows a gentle positive slope of variability ( $\bar{M} = 4.3$  per cent) with significance  $\nabla_S = 6.3$  over the entire time-scale of observations (see Fig. 4). Cui et al. (2010) carried out 5 GHz monitoring over a period of several years and found no significant variability. This object has been monitored for 15 yr at 408 MHz (Bondi et al. 1996) with a modulation index of 3.2 per cent. Also at 408 MHz, over a period of 5 yr Fanti et al. (1981) report a fractional variability  $\Delta S/S = 10$  per cent.

This is a well-known peaked-spectrum source at low frequencies (see also Callingham et al. 2017). It is detected in the NVSS survey

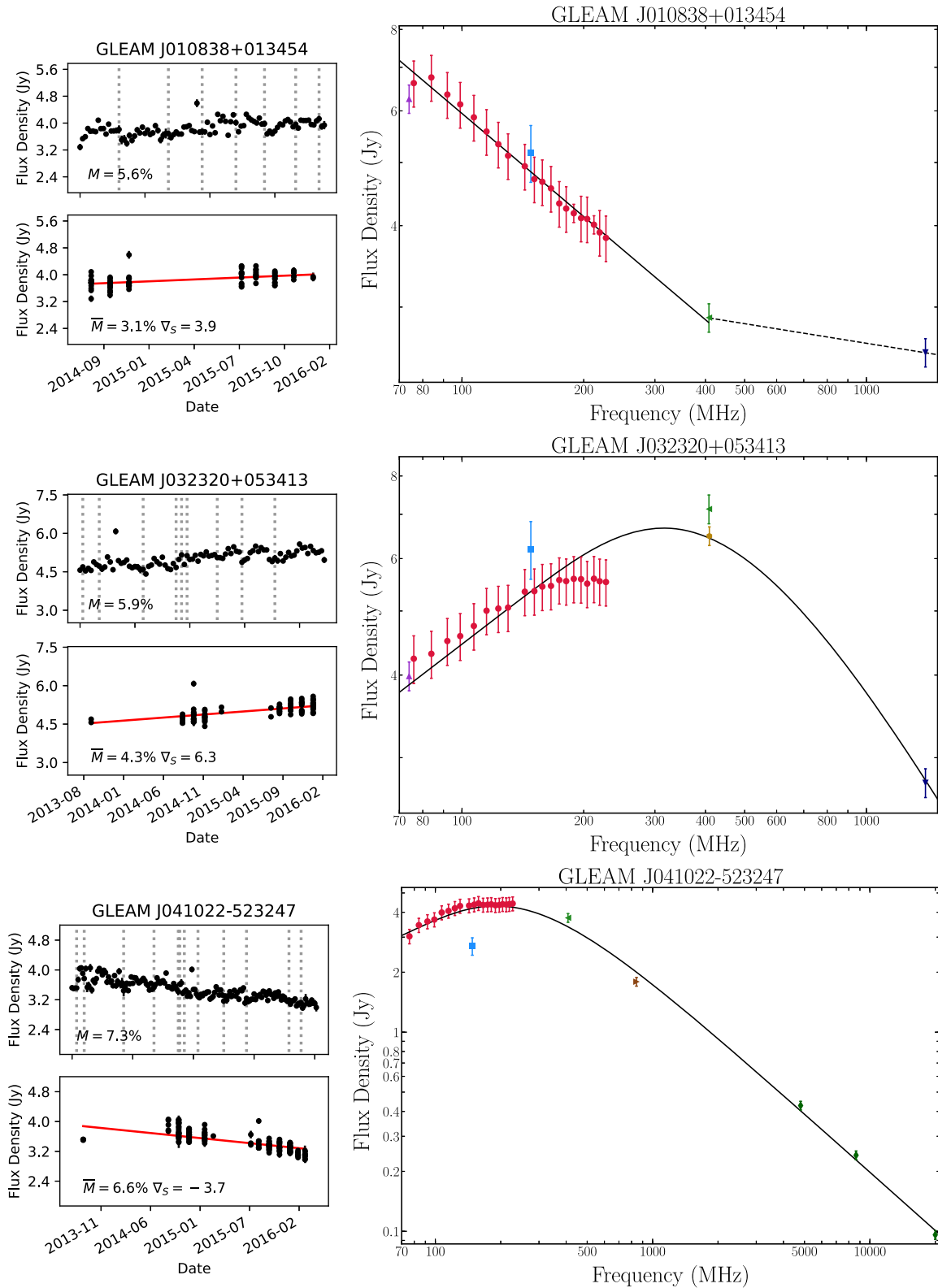
and the pronounced nature of the peak in the spectrum is quite sharp using our fitting technique. We have included in the SED the averaged data point of  $6.49 \pm 0.21$  Jy reported by Bondi et al. (1996), which is shown by an orange hexagon.

#### 4.3.3 GLEAM J041022–523247 (PKS B0409–526)

GLEAM J041022–523247 has an averaged modulation index of  $\bar{M} = 6.6$  per cent and has a peaked spectrum at low frequencies (see also Callingham et al. 2017). There is an AT20G counterpart but no published VLBI measurements to confirm the compact nature of the radio source. This source has a P-quad = 0.99 meaning that a second-order polynomial fit is preferred. We note, however, that the polynomial fit is most likely biased by a small number of data points around 2013 September. The TGSS flux density measurement is significantly lower than the GLEAM EGC measurements; this could be a result of variability due to the time difference between the two surveys (TGSS was observed between 2010 and 2012).

#### 4.3.4 GLEAM J051011–121042 (PKS B0507–122)

GLEAM J051011–121042 shows a negative slope of variability with significance  $\nabla_S = -5.1$ . The spectral index is best described as a power law. There are no counterparts in the AT20G or ICRF catalogues, and no reported redshift in the literature but there is a faint galaxy observable in the SuperCOSMOS Sky Survey images (Peacock et al. 2016) of the field. The NVSS image shows an



**Figure 4.** Light curves and SEDs for candidate variables identified in MWATS. The upper panel of each light curve shows the flux density as a function of sequential epoch number. The dashed grey line denotes an epoch number at which the time difference to the previous observation was greater than 8 d. The lower panel shows the flux density as a function of calendar date. The red line shows the linear fit to the averaged data points. The variability statistics for the raw ( $M$ ) and time-averaged ( $\bar{M}$ ,  $\nabla_S$ ) light curves are given in the appropriate panel. The SEDs are shown to the right; the red data points show the 20 point low-frequency spectrum measured from the GLEAM EGC. Archival data points are also shown and described in Section 4.2.



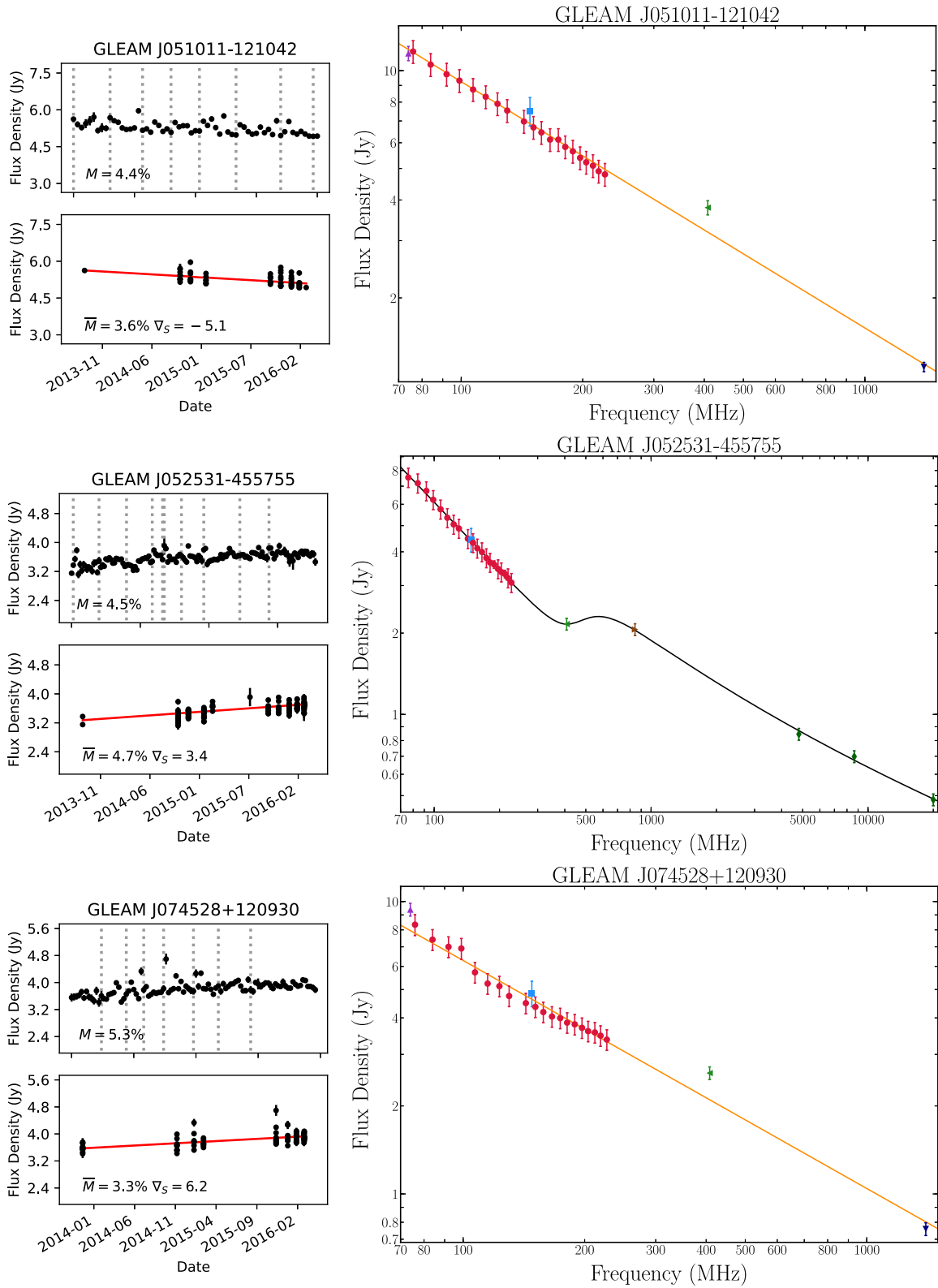


Figure 5. Continued: Light curves and SEDs for variables. Details as in Fig. 4.

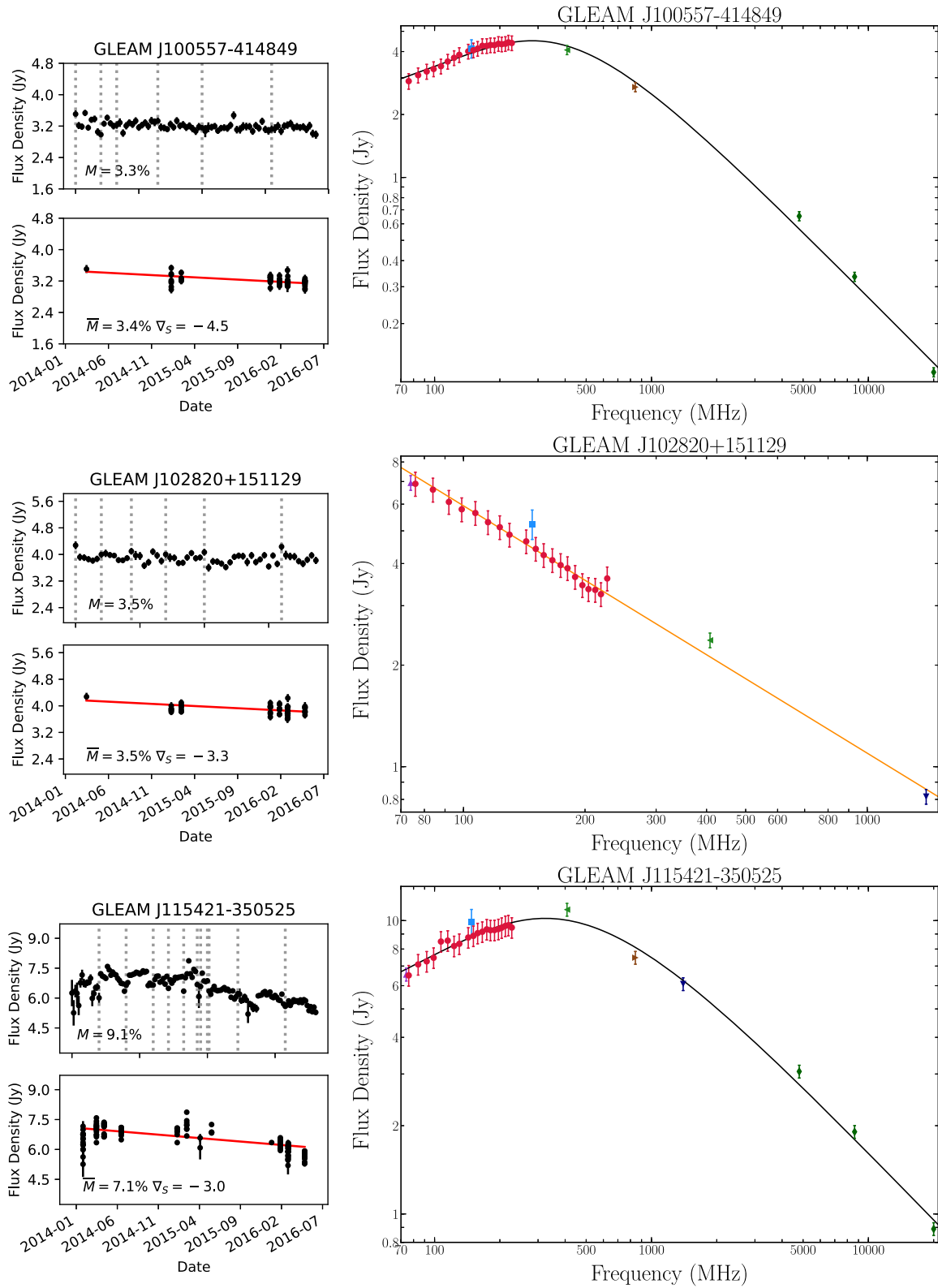


Figure 6. Continued: Light curves and SEDs for variables. Details as in Fig. 4.

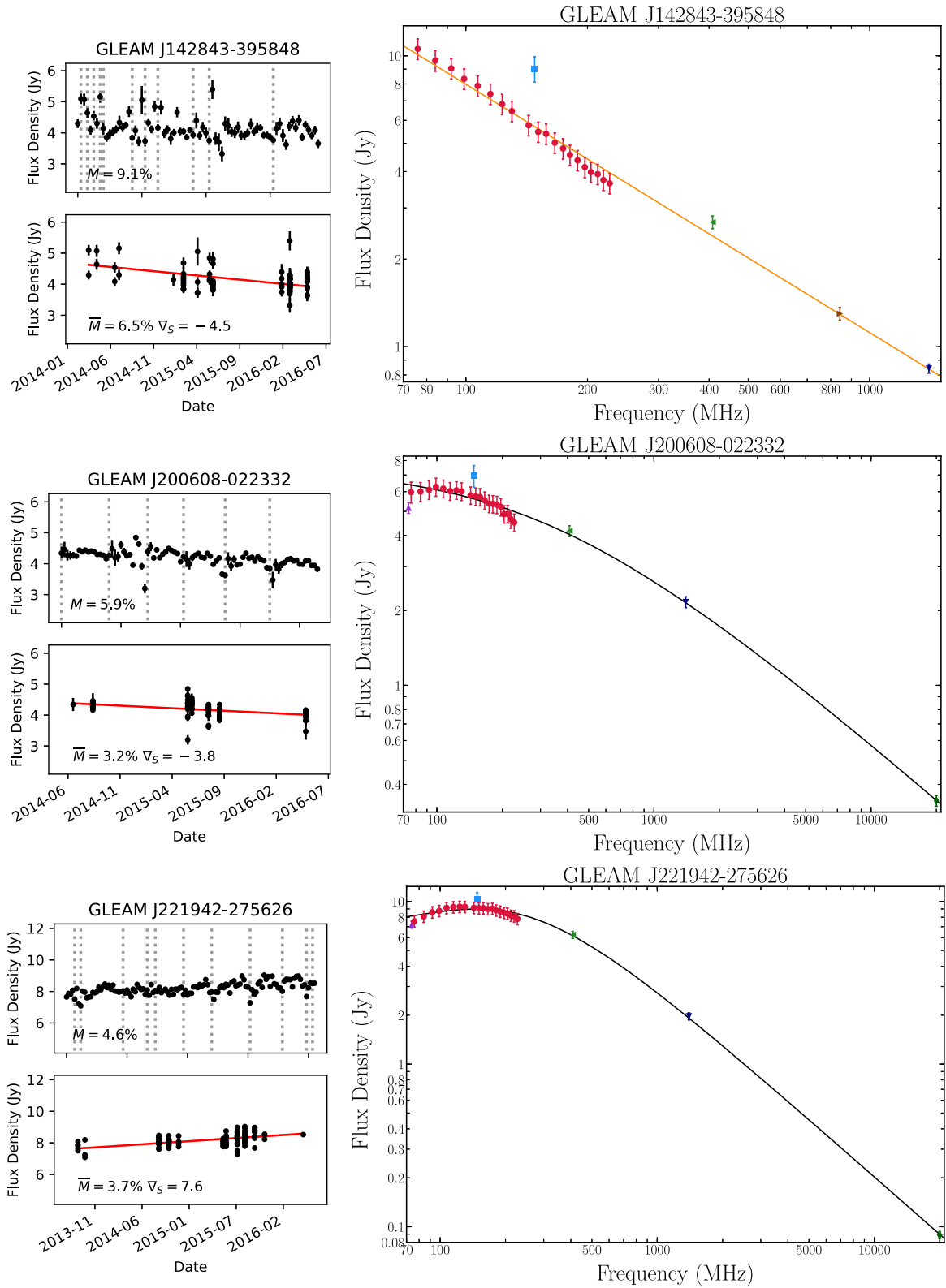
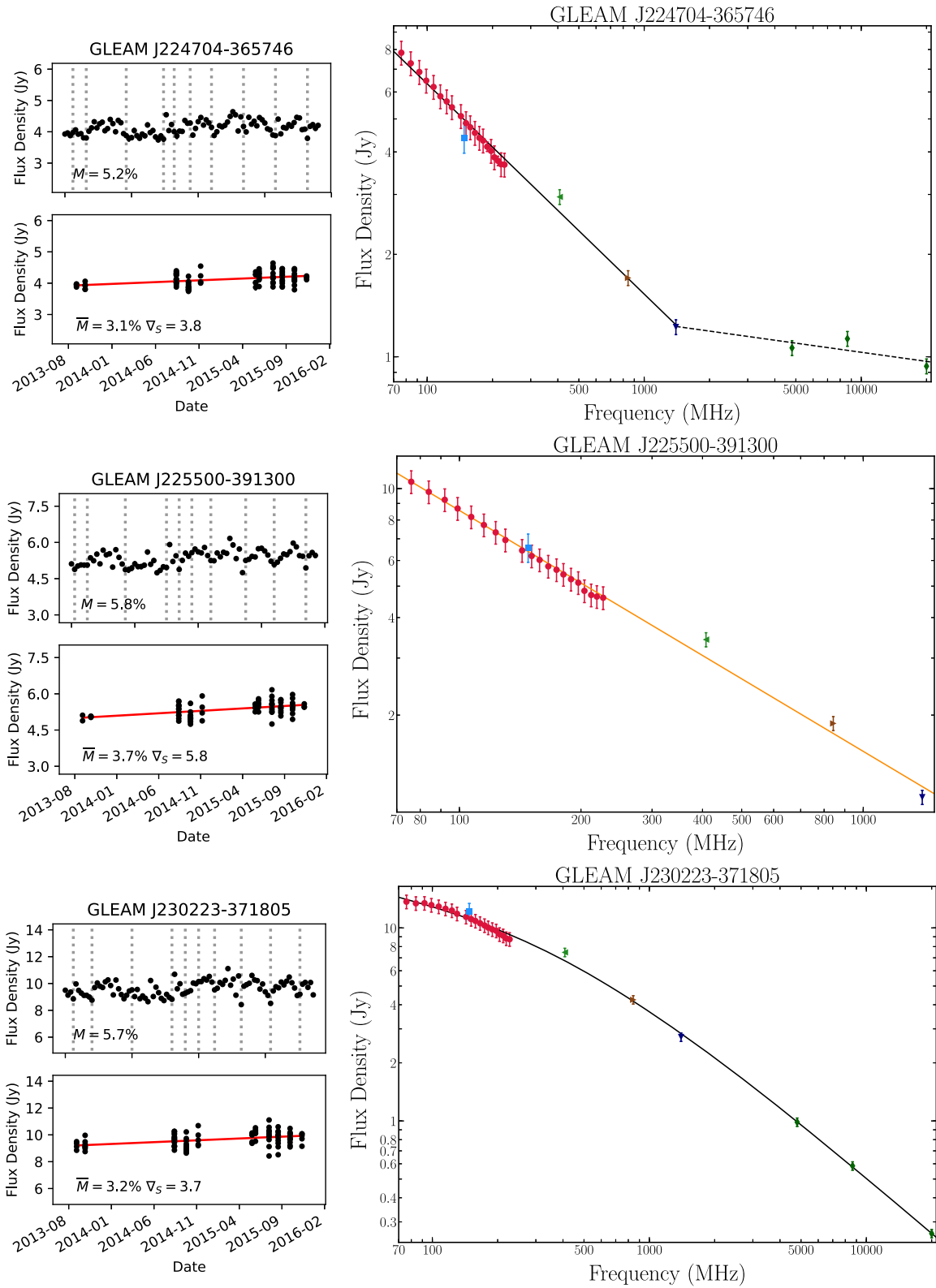


Figure 7. Continued: Light curves and SEDs for variables. Details as in Fig. 4.



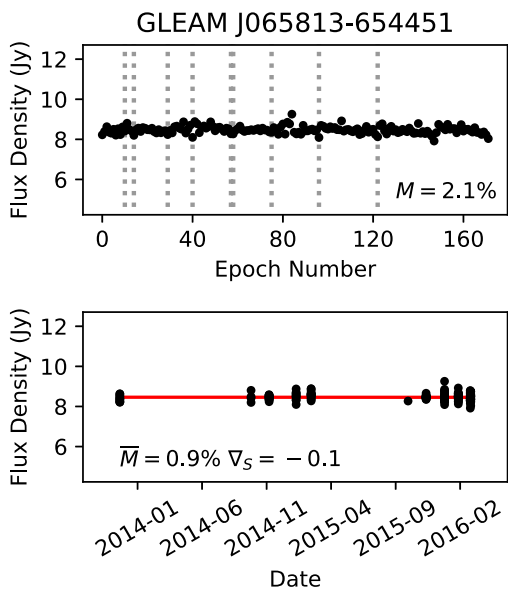
**Figure 8.** Continued: Light curves and SEDs for variables. Details as in Fig. 4.

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**Table 2.** Table summarizing properties of the 15 candidate variables. The  $\alpha$ -Fit column gives the best spectral classification with ‘p-law’ denoting a power law, ‘Broken’ denoting a two-component model and ‘Peak’ denoting a peaked spectrum; see Section 4.3 for full details. The class column gives the source class where ‘QSO’ denotes a Quasar, ‘G’ denotes galaxy, ‘BL’ a Blazar, and ‘BF’ denotes a blank field where no optical identification has been made. The TGSS column gives the morphology seen in the TGSS images, where ‘Pt’ denotes point-like, ‘SI Ext’ denotes slightly extended, and ‘Ext’ denotes a clearly extended source.

Name	$z$	$M$ (per cent)	$\overline{M}$ (per cent)	$\nabla_S$	P-slope	P-quad	$\alpha$ -Fit	Extent (mas)	Class	TGSS
GLEAM J010838+013453	2.107	$5.6 \pm 0.5$	$3.1 \pm 0.3$	3.9	0.94	0.64	Broken	<0.2	QSO/BL	SI Ext
GLEAM J032320+053413	0.1785	$5.9 \pm 0.7$	$4.3 \pm 0.5$	6.3	0.99	0.63	Peak	<200	G	Pt
GLEAM J041022–523247	–	$7.3 \pm 0.9$	$6.6 \pm 0.5$	–3.7	0.94	0.99	Peak	–	BF	Pt
GLEAM J051011–121042	–	$4.4 \pm 0.6$	$3.6 \pm 0.5$	–5.1	0.98	0.51	p-law	–	G	Ext
GLEAM J052531–455755	1.479	$4.5 \pm 0.7$	$4.7 \pm 0.4$	3.4	0.90	0.53	Broken	<10	QSO/BL	Pt
GLEAM J074528+120930	–	$5.3 \pm 0.5$	$3.3 \pm 0.3$	6.2	0.99	0.50	p-law	–	BF	Pt
GLEAM J100557–414849	0.81:	$3.3 \pm 0.7$	$3.4 \pm 0.4$	–4.5	0.97	0.77	Peak	–	G	Pt
GLEAM J102820+151129	–	$3.5 \pm 0.7$	$3.5 \pm 0.5$	–3.3	0.92	0.95	p-law	–	BF	Pt
GLEAM J115421–350525	0.258	$9.1 \pm 0.8$	$7.1 \pm 0.6$	–3.0	0.85	0.89	Peak	<15	G	Pt
GLEAM J142843–395848	–	$9.1 \pm 1.1$	$6.5 \pm 0.7$	–4.5	0.97	0.77	p-law	–	G	Ext
GLEAM J200608–022332	1.457	$5.9 \pm 0.5$	$3.2 \pm 0.3$	–3.8	0.94	0.60	Peak	–	QSO	Ext
GLEAM J221942–275626	0.657	$4.6 \pm 0.4$	$3.7 \pm 0.3$	7.6	0.99	0.50	Peak	–	G	Pt
GLEAM J224704–365746	2.25	$5.2 \pm 0.5$	$3.1 \pm 0.3$	3.8	0.92	0.53	Broken	<0.44	QSO	Pt
GLEAM J225500–391300	0.2615	$5.8 \pm 0.6$	$3.7 \pm 0.4$	5.8	0.98	0.50	p-law	–	G	SI Ext
GLEAM J230223–371805	1.140	$5.7 \pm 0.6$	$3.2 \pm 0.3$	3.7	0.92	0.5	Peak	–	G	Pt



**Figure 9.** Example of a non-variable source.

extended source (with a size 35 arcsec) and it is clearly extended in the TGSS image.

#### 4.3.5 GLEAM J052531–455755 (PKS B0524–460)

GLEAM J052531–455755 is classified as a Quasar/Blazar (Massaro et al. 2009) at redshift  $z = 1.479$  (Stickel, Kuehr & Fried 1993), and has been detected as a Fermi gamma-ray source (Acero et al. 2015). The spectrum is best described by a steep-spectrum power-law component with a break around 400 MHz, and then continued power-law component at higher frequencies. We note our fitting technique produces a quite complex fit, which is suggestive of a spectral peak between 500 and 1000 MHz and a steep spectrum at lower frequencies. It has a counterpart in AT20G and

both VSOP and ICRF catalogues. VLBI observations at 8 GHz show an unresolved compact object with a size <10 mas (Fey & Charlot 2000).

#### 4.3.6 GLEAM J074528+120930 (PKS B0742+122)

GLEAM J074528+120930 has no published optical counterpart. The radio source shows a positive slope of variability with significance  $\nabla_S = 6.2$ . The spectral index is well described by a simple power law. There is no known ICRF counterpart.

#### 4.3.7 GLEAM J100557 – 414849 (PKS B1003 – 415)

Burgess & Hunstead (2006) identify this source with a galaxy at an estimated photometric redshift of  $z = 0.81$ . It has a spectral turnover at  $\sim 300$  MHz and average modulation index  $\overline{M} = 3.4$  per cent. It has an AT20G detection but no published VLBI measurements.

#### 4.3.8 GLEAM J102820 + 151129 (PKS B1025 + 154)

GLEAM J102820+151129 has no published optical counterpart. The source has a VLSS measurement, no reported VLBI measurements and the spectrum is well fitted by a power law. The MWA spectrum falls below that expected from the MRC and TGSS data points; this could potentially be a result of the difficulty in calibrating the absolute flux scales between these surveys.

#### 4.3.9 GLEAM J115421–350525 (PKS B1151 – 348)

GLEAM J115421–350525 is identified with a broad-line radio galaxy at redshift  $z = 0.258$ , and is part of the well-studied 2 Jy sample of bright radio sources (Morganti, Killeen & Tadhunter 1993; Tadhunter et al. 1993). The source has the highest averaged modulation index of  $\overline{M} = 7.1$  per cent within our sample. The radio spectrum is peaked, with a turnover at  $\sim 400$  MHz. There is a measurement in AT20G and it is an ICRF source. Unpublished

maps<sup>2</sup> show that at 2.3 GHz the core is contained within 15 mas. Gaensler & Hunstead (2000) report variability with a modulation index of 1.9 per cent at 843 MHz.

#### 4.3.10 GLEAM J142843–395848 (PKS B1425–397)

This object has an optical identification in the SuperCOSMOS Sky Survey images (Peacock et al. 2016) but no redshift is reported in the literature. There are no AT20G or VLBI measurements. The MWA spectrum is well fitted with a power law, but archival data points suggest a steeper power-law spectrum. This source is near ( $\sim 12^\circ$ ) to the powerful radio galaxy Centaurus A, and due to its brightness it can be complex to extract measurements from this region; this could explain the discrepancy here. The source is strongly extended in the TGSS image and is resolved into two components in the Australia Telescope-PMN catalogue with 5 GHz flux densities 115 and 128 mJy, respectively (McConnell et al. 2012).

#### 4.3.11 GLEAM J200608–022332 (PKS B2003 – 025)

This object is identified with a QSO at redshift  $z = 1.457$  (Aldcroft, Bechtold & Elvis 1994). There are VLSS and AT20G counterparts, but no VLBI measurements. Callingham et al. (2017) identify this as a peaked-spectrum source. It showed low levels of variability ( $\bar{M} = 3.2$  per cent) in this study. This source is extended in the TGSS image and unresolved in NVSS.

#### 4.3.12 GLEAM J221942–275626 (PKS B2216–281)

GLEAM J221942–275626 is a galaxy at redshift  $z = 0.657$  (McCarthy et al. 1996). The source shows a steady increase in flux density with significance  $\nabla_S = 7.6$  and a low-frequency turnover in the spectrum with a peak around 150 MHz. There is a counterpart in the AT20G catalogue but not in ICRF.

#### 4.3.13 GLEAM J224704–365746 (PKS B2244 – 372)

GLEAM J224704–365746 is a QSO at redshift  $z = 2.25$  (Wilkes et al. 1983). The radio spectrum is well fitted with a steep power law for frequencies below 1 GHz. At higher frequencies the data are better fitted with a flatter spectrum (shown by the dashed line). The source is detected in AT20G, and VLBA observations show that it is compact with an angular diameter  $< 0.44$  mas (Petrov et al. 2006).

#### 4.3.14 GLEAM J225500–391300 (PKS B2252–394)

GLEAM J225500–391300 is identified with a  $z = 0.2615$  galaxy, observed as part of the 6dF Galaxy Survey (Jones et al. 2009). The source shows a steady positive slope in flux density with significance  $\nabla_S = 5.8$ . The radio spectrum is well described by a simple power law. There is no counterpart in the ICRF catalogue but there is a counterpart in the AT20G and ATPMN catalogues. This source is slightly extended in the NVSS and TGSS images.

#### 4.3.15 GLEAM J230223–371805 (PKS B2259–375)

This object has a peak in the SED with a low-frequency turnover below 100 MHz (Callingham et al. 2017). It is associated with a redshift  $z = 1.140$  galaxy (Mahony et al. 2011) and it is a point source in both the TGSS and NVSS images. It has counterparts in the AT20G and PMN surveys. It is listed as a calibrator in the VLBA data base (Petrov et al. 2006) but no information about the angular size is given.

## 5 DISCUSSION

Summarizing the information in Section 4.3: of the 15 sources, five are well described by a power-law spectral index; seven have a peak and turnover at MHz frequencies and three are not well fitted by a single spectral component. Five of the 15 sources have published VLBI measurements confirming the presence of compact components between 0.2 and 200 mas at frequencies  $> 2$  GHz. Further sources have unpublished VLBI measurements<sup>3</sup> but lack sufficient information to characterize source size. The optical properties of this sample are quite diverse: it contains four QSOs (two of which are also classed as Blazars), eight galaxies (six of which have reported redshifts), and three fields where no optical identification has been made.

Of our total sample, 1.6 per cent (15/944) are classed as possible variables. There are 54 peaked-spectrum sources, which accounts for only 5.7 per cent of the total, yet seven sources (47 per cent) in our variable sample have a peaked spectrum. There is therefore a significant excess in the occurrence of variability amongst peak spectrum sources. This is consistent with the findings of Chhetri et al. (2018), who found that every peaked-spectrum source was compact on subarcsec scales and that these peaked sources dominated the compact source population. As peaked-spectrum sources are often young and compact there is an increased probability for variability via refractive scintillation.

Five objects show varying degrees of extended radio emission (see Table 2). One of the extended objects (GLEAM J010838+013453) has VLBI data at 5 GHz constraining the size to less than 0.2 mas. Variability in extended or resolved sources would not necessarily be predicted. However, it is difficult to completely disentangle the geometry of these objects as well as episodic periods of past radio activity. For example, GLEAM J010838+013453 could have an extended radio structure from older radio activity that has expanded significantly, while more recent activity remains compact and unresolved. We discuss this further below with respect to the possible mechanisms for variability.

The minimum averaged modulation  $\bar{M}$  index we measure is 3.1 per cent and the maximum is 7.1 per cent. We emphasize that these sources display the largest amount of long-term variability (based on our definition) within this sample. In comparison with historical surveys, where highly variable objects are often classed as having modulation indices of  $> 50$  per cent, these objects would possibly be considered non-variable (or certainly low-level variables).

Our selection method is biased towards detecting long-term variability, so it is possible that  $\nabla_S$  could remain below 3.0 but short-term, high modulation index variability could still be present. Independent inspection of high  $M$  and  $\bar{M}$  outliers has shown they were due to erroneous data of instrumental origin, although not every data

<sup>2</sup><http://astrogeo.org/rfc/>

<sup>3</sup>[http://astrogeo.org/vlbi/solutions/rfc\\_2017d/rfc\\_2017d\\_cat.txt](http://astrogeo.org/vlbi/solutions/rfc_2017d/rfc_2017d_cat.txt)

point could be checked by hand. We therefore cannot rule out the possibility that some of the low modulation, short-term variability could be of astrophysical origin. We can say, however, that sources with a large modulation index ( $M$  and  $\bar{M} > 10$  per cent) arising from short-term variability are rare. Future work is directed at applying more advanced statistical techniques to identify short-term variability over the entire MWATS sample.

Before we discuss the most probable mechanism of variability within this sample, it is useful to explore theoretical predictions of the time-scale and amplitude expected for the cases of refractive scintillation and intrinsic jet-driven variability. We note that the debate surrounding the intrinsic versus extrinsic nature of variability in low-frequency sources has largely been concluded (see Rickett 1986). It is, however, pertinent to frame the results in this paper not only with respect to previous work but also in reference to upcoming surveys with next generation instruments such as the SKA.

### 5.1 Intrinsic variability

If we assume that the variability observed in our sample is driven solely by the synchrotron process and that the Compton catastrophe limit holds (Kellermann & Pauliny-Toth 1969), then the maximum brightness temperature we can achieve is  $T_b < 10^{12}$  K (assuming there is no beaming). We can therefore define the following expression (Miller-Jones et al. 2008):

$$T_b = \frac{S_\nu c^2}{2k_B \nu^2 \Omega} \leq 10^{12} \text{ K}, \quad (2)$$

where  $S_\nu$  is the amplitude of variability,  $k_B$  is the Boltzmann constant,  $\nu$  is the observing frequency, and  $\Omega$  is the solid angle subtended by the source. The solid angle can be approximated by  $\Omega = r^2/D^2$ , where  $r$  is the source size and  $D$  is the distance. We assume that  $r$  and thus the angle subtended by the source is limited by the light traveltime (or time-scale for variability)  $\tau$ , i.e.  $r = \tau c$ . By substituting in  $\Omega = D^2/\tau^2 c^2$ , we can then define the following expression, which gives the maximum flux density change permitted for a synchrotron emitting source:

$$\Delta S \leq \frac{2k_B \nu^2 \tau^2 T_b}{D^2}. \quad (3)$$

If we take  $\tau = 2.8$  yr (i.e. the maximum time-scale sampled in this survey) and  $D = 10$  Gpc (a reasonable assumption given our sample contains objects with known distances between  $\sim 1$  and 15 Gpc), we find  $\Delta S < 5.3 \times 10^{-4}$  Jy.

In reality, we measure changes that are much greater, i.e.  $\Delta S \sim 0.5$  Jy; see the left-hand panel of Fig. 10 for a plot of the expected brightness temperature as a function of amplitude of variability ( $\Delta S$ ). Assuming  $\Delta S \sim 0.5$  Jy and recalculating the permitted time-scales of variability, we find  $\tau > 100$  yr. See the right-hand panel of Fig. 10 for a plot of variability time-scale  $\tau$  as a function of amplitude of variability ( $\Delta S$ ). To obtain brightness temperatures below  $10^{12}$  K the amplitude of variability is therefore quickly quenched and varies as  $T_b \propto \tau^{-2}$ .

Relativistic beaming and Doppler boosting can increase the maximum brightness temperature permitted in these scenarios and scales as  $b^3 T_b$  (see Lähteenmäki & Valtaoja 1999). Assuming a boosting factor of  $b = 10$  and recalculating the example above, we find allowed flux density variations to be as large as  $\Delta S = 0.53$  Jy, with time-scales of  $\tau > 3.16$  yr. This is much more in line with our results. Two of our variables are classed as Blazars and it is plausible that some level of beaming is taking place and that this is contributing to the changes in flux density. We note that these objects are

also compact and have angular diameters less than 0.2 and 10 mas, respectively (see Table 2). In the next section, we will discuss how the compact nature of these objects can give rise to variations via refractive scintillation.

Assuming the absence of beaming, then the calculations above provide a strict limit on the permitted variability time-scales and flux density changes at low frequencies. As relativistic beaming is rare, it is unlikely that simple synchrotron radiation can explain the type of variability seen in our survey, as the variability time-scales are very long (as restricted by the Compton catastrophe limit) and are not well sampled by our survey.

### 5.2 Refractive interstellar scintillation

The time-scale for refractive scintillation is a function of wavelength and scales as  $\nu^{-\beta}$ , where the spectral index  $\beta$  is 11/3 for Kolmogorov turbulence (Armstrong, Rickett & Spangler 1995). Assuming at 5 GHz we observe scintillation time-scales as fast as 1 d, at 154 MHz we would expect these time-scales to be 5.8 yr. The modulation index of variability scales with frequency as  $(\nu/\nu_0)^{17/30}$ , where  $\nu_0$  is the transition frequency and the exponent arises from assuming Kolmogorov turbulence (Walker 1998).

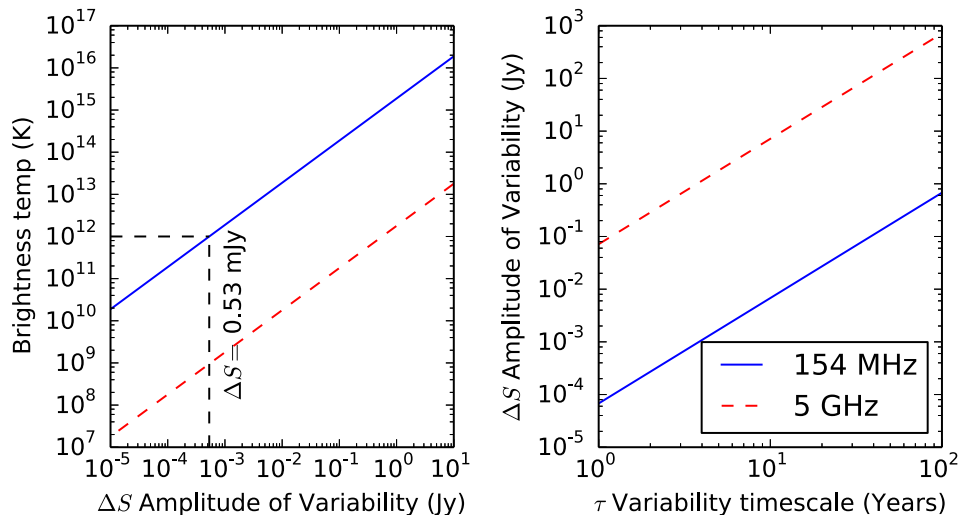
The transition frequency separates weak and strong scattering, with weak scattering dominating the higher frequencies. The transition frequency is a function of Galactic latitude and is as high as 40 GHz within the Galactic plane and as low as  $\sim 6$  GHz around the Galactic poles (Walker 1998, Walker 2001). Assuming a modulation index of 100 per cent at the transition frequency, at 154 MHz we would predict a modulation index of 12.7 per cent at the Galactic pole (with  $\nu_0 = 6$  GHz) and 4.4 per cent towards the Galactic centre (with  $\nu_0 = 40$  GHz). These values are consistent with our results (i.e. Median  $\bar{M} = 2.1 \pm 2.0$  per cent).

At low frequencies the radio spectra of compact sources can often include extended steep-spectrum relic emission from earlier periods of source activity. When combined with generally larger beam sizes, the contribution from compact flat-spectrum components, which may be prominent at higher frequencies, will therefore be strongly diluted. These factors will play an important role in the prevalence and detectability of low-frequency radio variability as a function of time, frequency, and angular resolution. For refractive scintillation, above a certain angular size cut-off angular broadening will take place and the scintillation and modulation index will be quenched. The quenching amount depends on the ratio of the angular size cut-off ( $\theta_{\text{lim}}$ ) and source size ( $\theta_s$ ) as  $(\theta_{\text{lim}}/\theta_s)^{7/6}$  (Walker 1998).

If we take the object GLEAM J010838 + 013454 as an example, VLBI observations show a core size of 0.2 mas at 5 GHz. Using the NE2001 electron density model<sup>4</sup> (Cordes & Lazio 2002), we can calculate the angular size cut-off (or angular broadening parameter) as a function of location on the sky. We find that for GLEAM J010838 + 016453, we require the source size to be less than, or the majority of the flux density to be contained within, 48.5 mas at 154 MHz (calculated with a transition frequency of 8.0 GHz). As VLBI measurements of the core show a size less than 0.2 mas, it clearly has a compact high-frequency core. This source is, however, slightly extended in the 150 MHz TGSS image with a deconvolved major axis of  $\approx 15$  arcsec.

The transition from a flat spectrum to a steep spectrum at  $\nu \approx 400$  MHz is shown clearly in the SED in Fig. 4. It is complicated to completely untangle the geometry here, especially as a function of

<sup>4</sup><https://www.nrl.navy.mil/rsd/RORF/ne2001/>



**Figure 10.** Left-hand panel: A plot of amplitude of variability ( $\Delta S$ ) versus brightness temperature for a hypothetical source. We assume in this example that the source is at a distance of 10 Gpc and the variability time-scale is 2.8 yr. To not break the brightness temperature limit of  $10^{12}$  K, the amplitude of variability ( $\Delta S$ ) is restricted to be below 0.53 mJy (at 154 MHz). Right-hand panel: A plot of variability time-scale ( $\tau$ ) versus amplitude of variability ( $\Delta S$ ). In this case, we assume that  $T_b = 10^{12}$  K and we find that for typical variations of  $\Delta S = 0.5$  Jy we require variability time-scales to be greater than  $\sim 100$  yr (at 154 MHz). The comparison case of 5 GHz is shown with the red dashed line.

time and frequency. For example, relic low-frequency radio emission may be present as well as more recent compact emission. In all cases, however, providing a significant fraction of the flux is contained within the angular size cut-off, refractive scintillation is still possible.

One final possibility not fully explored here is that of variability caused by a free-free absorbing screen of material moving across the line of sight. This has been shown to cause significant spectral variability in the peaked-spectrum source PKS 1718 – 649 and would possibly alleviate the need for all flux to be contained within a compact region (see Tingay et al. 2015, for a full discussion).

Both the time-scale and expected modulation for refractive scintillation are in line with our results and are the plausible explanation for variability within our sample, although we cannot completely rule out a contribution from beamed synchrotron emission. In both the synchrotron (non-beamed) and refractive scintillation cases the time-scale of variability is typically longer than our survey length and justifies why fitting the gradient of the light curves is a reasonable strategy for detecting and characterizing variability.

### 5.3 Comparison with previous surveys

The original Hunstead (1972) paper reported four variables: PKS B1504 – 166, PKS B1524 – 136, PKS B2230 + 114, and PKS B2251 + 158 at 408 MHz. The source PKS B2251 + 158 is located at the edge of our survey range and we did not have many measurements for this object. The remaining three sources are plotted in Fig. 11. All three objects are classified as non-variables using the criteria we define in this paper.

The source PKS B1504 – 166 has the highest modulation index in archival observations ( $M = 16.7$  per cent) whereas we measured only  $\bar{M} = 2.3$  per cent. VLBI observations of PKS B1524 – 136 at 8 GHz by Mantovani et al. (2002) show a double-sided jet with angular extent up to 300 milliarcsec. Observations by Jeyakumar et al. (2000) show that 50 per cent of the total flux density is from a compact structure with a size of  $< 100$  milliarcsec. It has been

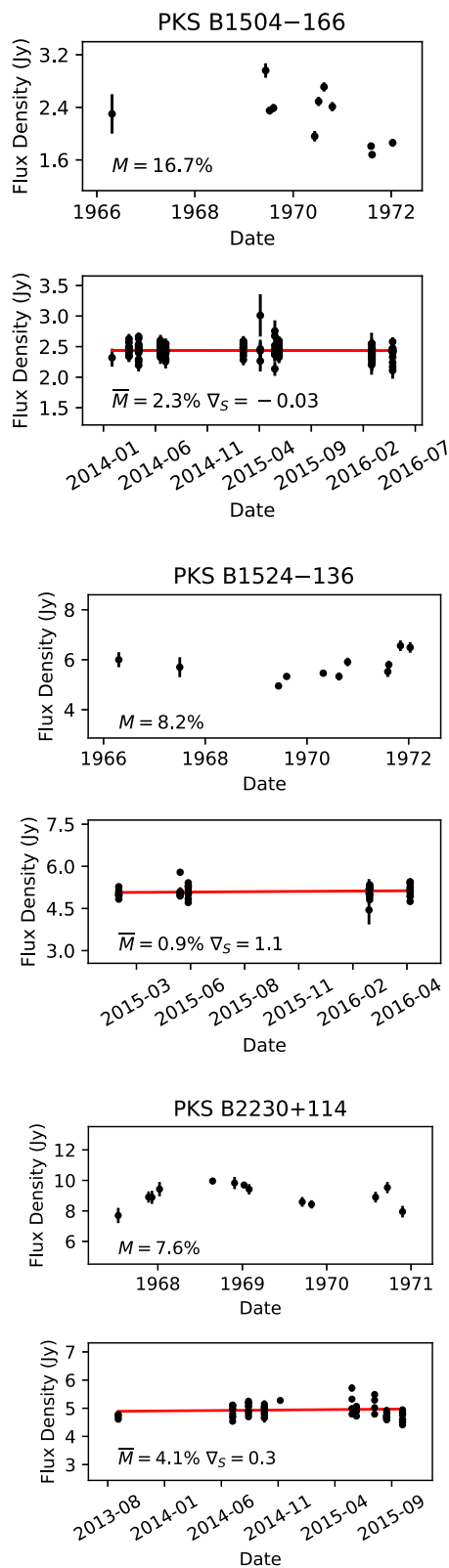
monitored at 408 MHz for 15 yr by Bondi et al. (1996), who report a modulation index of 8.2 per cent.

The source PKS B2230 + 114 has also been studied at 408 MHz for 15 yr by Bondi et al. (1996). We only measure an  $\bar{M} = 4.1$  per cent, whereas Hunstead (1972) measured 7.6 per cent and Bondi et al. (1996) report a value of 4.3 per cent. VLBI observations by Xu, Wehrle & Marscher (1998) constrain the core size to  $< 2$  milliarcsec. The angular size limit at 154 MHz in this direction is 48 mas and the transition frequency is 9.31 GHz (calculated using NE2001; Cordes & Lazio 2002). Using the transition frequency and assuming 100 per cent modulation at this transition, we would predict modulation indices of 9.8 and 17.0 per cent at 154 and 408 MHz, respectively. In both the 154 and 408 MHz cases here we measure lower variability than predicted, but the ratios of our measurements and those of Hunstead (1972) are consistent.

Clearly, the variability of these three objects is being quenched at lower frequencies. Both PKS B2230+114 and PKS B1524–136 have peaked spectra, while PKS B1504–166 has a flatter spectrum, all of which are indicators of compactness. As discussed above, the amount of LFV is largely a function of frequency and source size. Potentially, the object PKS B1524 – 136 has become extended such that refractive scintillation has become quenched at 154 MHz, or, that the frequency shift has reduced the modulation and lengthened the time-scales. What is interesting about our sample is that we have variables that have larger modulation indices than the equivalent 154 MHz measurements of the Hunstead (1972) variables. It would therefore be interesting to obtain higher frequency time-domain measurements of our sample, say at 400 MHz, to determine if the variability amplitudes are indeed larger at lower frequencies or if this is due to intermittency of source variability.

The total number of variables in our sample is 15/944 or 1.6 per cent. These numbers are consistent with McGilchrist & Riley (1990), who show that 1.1 per cent of a sample of 811 sources show fractional variability  $> 15$  per cent on time-scales of 1 yr. Although it is difficult to compare fractional variability versus modulation index, we find no variables with a modulation index greater than 15 per cent. We also find a lower abundance of variability when





**Figure 11.** Upper panel: Light curves of PKS B1504–166, PKS B1524–136, PKS B2230+114; measurements reported in Hunstead (1972) at 408 MHz. Lower panel: MWA measurements from this survey.

compared with intermediate frequencies (e.g. 408 MHz) where 25 percent of compact sources and 51 percent of flat spectrum sources display dynamic properties (McAdam 1979; Fanti et al. 1981; Spangler et al. 1989). All of the surveys above have flux density limits greater than 1 Jy.

## 6 CONCLUSIONS

In this paper, we report the detection of 15 candidate low-frequency variables using time-domain measurements from the MWA. The type of variability detected in this survey can be characterized as long duration ( $>5$  yr) with low modulation indices ( $<10$  per cent). In general, the variable low-frequency sky is very stable with only 1.6 percent of objects displaying any variability on time-scales of years. Of our variables a large proportion (7/15) have a peaked spectrum and we conclude that LFV is more prevalent in this spectral class. If we scale up the results from this sample to our entire survey size, i.e. approximately 350 000 radio sources, we could expect an upper limit of  $\sim 6000$  low-frequency variables in the Southern hemisphere.

We review the theory and conclude that at 154 MHz, synchrotron radiation (or intrinsic AGN variability) produces variations on time-scales of  $\sim 100$  yr and refractive scintillation produces variations on time-scales greater than 5 yr. The amount of variability decreases as a function of frequency, starting at 100 per cent at the transition frequency (typically around 5 GHz) to approximately 10 per cent at 150 MHz. Quite extreme relativistic and unrealistic boosting factors are needed to explain our results with respect to intrinsic synchrotron variations. However, we cannot completely rule this out as having an impact on the variations observed, nor can we totally rule out contributions from other models such as free-free absorption.

Based on our results, the long-time-scale time-variable sky will therefore be dominated by extragalactic sources that have angular diameters approximately less than 50 mas at 154 MHz, which are fairly rare. The majority of objects observed by SKA-Low ( $<200$  MHz) will therefore be non-variable. The short-time-scale (minutes to days) time-variable SKA-Low sky will be dominated by objects whose angular size permits diffractive scintillation, i.e. pulsars.

Coherent emission mechanisms and bright nearby synchrotron explosions will start to be detectable once the detection threshold drops below 1 mJy. It remains the subject of future work to thoroughly probe the short-time-scale parameter space with the data collected in this survey. For example, we are still refining the model of the primary beam, which is one of the largest uncertainties in our variability analysis.

The low-frequency radio sky is relatively inactive and offers a unique space to search for the counterparts to gravitational wave events. The error box currently produced by the Laser Interferometric Gravitational Wave Observatory (LIGO; Aasi et al. 2015) can be up to a few thousand square degrees on the sky. This can efficiently be covered with a few MWA pointings (Kaplan et al. 2016). The MWA is therefore an effective instrument to search for electromagnetic signatures of merger events, not only because the field of view is so large, but also because the background rate of variable sources is low.

At GHz frequencies the background abundance of highly variable radio sources is much higher than at low frequencies, i.e. there are approximately  $3 \text{ deg}^{-2}$  (or 3000 per 1000  $\text{deg}^2$ ) at 5 GHz that show modulation indices over 50 percent on month time-scales (Bell et al. 2015), which is the possible rise time for a merger

event afterglow. Currently in this survey we detect no variables with such a high modulation index, therefore on a time-scale of months the background low-frequency sky will be static. This will be advantageous for identifying gravitational wave counterparts.

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