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Carbon Abundances in Starburst Galaxies of the Local Universe

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Abstract

The cosmological origin of carbon, the fourth most abundant element in the universe, is not well known and a matter of heavy debate. We investigate the behavior of C/O to O/H in order to constrain the production mechanism of carbon. We measured emission-line intensities in the spectral range from 1600 to 10000 Å on Space Telescope Imaging Spectrograph (STIS) long-slit spectra of 18 starburst galaxies in the local universe. We determined chemical abundances through traditional nebular analysis, and we used a Markov Chain Monte Carlo method to determine where our carbon and oxygen abundances lie in the parameter space. We conclude that our C and O abundance measurements are sensible. We analyzed the behavior of our sample in the [C/O] versus [O/H] diagram with respect to other objects such as DLAs, neutral ISM measurements, and disk and halo stars, finding that each type of object seems to be located in a specific region of the diagram. Our sample shows a steeper C/O versus O/H slope with respect to other samples, suggesting that massive stars contribute more to the production of C than N at higher metallicities, only for objects where massive stars are numerous; otherwise, intermediate-mass stars dominate the C and N production.

Key words: galaxies: evolution – galaxies: starburst – Galaxy: abundances – H II regions – ultraviolet: galaxies

Supporting material: figure sets

1. Introduction

Carbon is the fourth most abundant element in the universe as well as one of the key ingredients for life as we know it. It is a ubiquitous element in the interstellar medium (ISM): most molecules in the ISM are C-bearers (with CO being the most abundant molecule), and carbonaceous dust particles represent an important fraction of the ISM dust composition (e.g., Garnett et al. 1995; Dwek 1998, 2005; Roman-Duval et al. 2014; Zhukovska 2014, and references therein). Carbon also has an important role in regulating the temperature of the ISM: it contributes to the heating of the interstellar gas because it is the main supplier of free electrons in diffuse clouds, and it also contributes to the cooling of the warm interstellar gas through the emission of 158 μm C II (Stacey et al. 1991; Gullberg et al. 2015). Despite its relevance, there are only a few nebular carbon abundance determination studies because its brightest collisionally excited lines (CELs) are [C III] 1907 and C III] 1909 Å⁵ and [C II] 2326 Å in the UV (e.g., Dufour et al. 1982; Garnett et al. 1995, 1999; Peña-Guerrero et al. 2012b; Stark et al. 2015a), and [C II] 158 μm in the far-IR (e.g., Tielens & Hollenbach 1985; Stacey et al. 1991; Madden et al. 1997; Garnett et al. 2004; Canning et al. 2015). The brightest carbon recombination line (RL), C II 4267 Å, is in the optical, and it is only detectable and measurable in bright Galactic and extragalactic nebulae (e.g., Esteban et al. 2002, 2005, 2004, 2014; García-Rojas et al. 2005; López-Sánchez et al. 2007; Peña-Guerrero et al. 2012b). RLs are intrinsically very weak, hence a large resolving power is required to accurately measure such lines.

It is generally agreed that massive stars ($M > 8 M_{\odot}$) synthesize most of the oxygen, whereas carbon is synthesized

by both low- and intermediate-mass stars as well as by massive stars (e.g., Clayton 1983; Cowley 1995; Henry et al. 2000). In low-metallicity environments, carbon is thought to be “primary.” Primary nucleosynthesis is defined as all of those nuclei that can be produced from the initial H and He present in the star—it is independent of the initial stellar metallicity (e.g., ¹²C, ¹⁶O, ²⁰Ne, ²⁴Mg, ²⁸Si); secondary nucleosynthesis is defined as those nuclei that can be produced using preexisting nuclei from previous stellar generations, hence it is dependent on the initial stellar metallicity (e.g., ¹⁷O, ¹⁸O) (Meyer et al. 2008; Lugaro et al. 2012). A particularly interesting element is ¹⁴N, which is produced by both primary and secondary nucleosyntheses (e.g., Vincenzo et al. 2016 and references therein), but this element will not be discussed in this paper. Secondary production is common to stars of all masses (Matteucci 1986). In high-metallicity environments, C, N, and O are synthesized during the CNO cycle as catalysts to produce He in both intermediate-mass and massive stars; after the hydrogen-burning phase (through the proton–proton chain or CNO cycle), carbon and oxygen are byproducts of the triple- α process (Renzini & Voli 1981). C and O are then brought to the surface of the star through the dredge-up process. As the mass of the star increases, He and C are removed from the star before forming O.

The debate on the mass of the stars that contribute the most to the production of carbon is complicated due to the existence of different yields in stellar chemical evolution models (Carigi et al. 2005). There are currently several uncertainties in the carbon yields of massive stars. If mass-loss rates depend on metallicity, then the yields of C and O also depend on metallicity, with C increasing at the expense of O (Garnett et al. 1995). The amount of carbon and oxygen ejected by a star is directly dependent on its mass and on its metallicity. Both ¹⁶O and ¹²C are products of the triple- α process. Stellar evolution

⁵ For simplicity, we will refer to [C III] 1907 and C III] 1909 Å as C III] 1907+09 Å.

models predict that the mass of the least massive star capable of producing and ejecting new oxygen is about $8 M_{\odot}$ (Garrett et al. 1995); less massive stars simply leave it in the core. In the case of carbon, the minimum stellar mass is predicted to be 2 to $8 M_{\odot}$ in order to account for both massive stars and intermediate-mass through the dredge-up process during their AGB phase (Boyer et al. 2013).

Previous works have shown that the yield of carbon varies with metallicity. Carigi et al. (2005) found that out of 11 Galactic chemical evolution models with different yields adopted for carbon, nitrogen, and oxygen, only two models fit the oxygen as well as the carbon gradient. These two models had carbon yields that increase with metallicity due to the winds of massive stars, and decrease with metallicity due to the winds of low- and intermediate-mass stars. Fabbian et al. (2009) determined [C/O] for 43 metal-poor halo stars, which are in reasonably good agreement with the results of Carigi et al. (2005). Boyer et al. (2013) suggested that a reduction in the carbon-rich to oxygen-rich AGB stars (respectively referred to as C and M stars) with metallicity is required in all modern TP-AGB models. This requirement comes from (i) the larger amount of carbon needed to be dredged up to make the $C/O > 1$ transition, and (ii) the third dredge-up starting later at higher luminosities and being less efficient at increasing metallicity. Note that a C star must have $C/O > 1$, whereas an M star has $C/O < 1$. When the effects of rotation are included in stellar models, there is a very large increase in the yields of primary C, N, and O at very low metallicity (Meynet & Maeder 2002). Furthermore, the majority of massive stars are suspected to be in close binaries and experience interaction with a companion (Sana et al. 2012). The implications for the yields of massive stars have not yet been systematically studied.

Carbon is crucial for the composition of interstellar dust. In order to obtain accurate ISM abundances, it is paramount to account for the presence of dust grains. Several studies have found that oxygen depletion in the Orion Nebula amounts to a correction in the total O/H of about 0.09 dex (Esteban et al. 1998, 2004; Mesa-Delgado et al. 2009; Simón-Díaz & Stasińska 2011). A correction of 0.09 to 0.11 dex was suggested by Peimbert & Peimbert (2010) for the O/H of H II regions, depending on the metallicity of the object. Most of the ISM dust can be broadly classified into either carbonaceous- or silicate-based, hence both types of dust affect our study of C/O. Dust formation can be broadly divided into two types of sources: (i) those that undergo quiescent mass loss (e.g., W-R stars) and (ii) those that return their ejecta eruptively back into the ISM (e.g., Type Ia and Type II supernovae and asymptotic giant branch [AGB] stars). The type of dust does not necessarily depend on the formation source but rather on the C/O in the ejecta. For low- and intermediate-mass stars ($M < 8 M_{\odot}$), if $C/O > 1$, all the oxygen is tied up in CO molecules and the newly formed dust grains will be carbon rich; if $C/O < 1$, the extra oxygen will combine with other elements to form silicate-based types of dust grains (Dwek 1998). Massive stars will contribute to the dust grain production only with the coolest stars and according to the exposed material (Cohen & Barlow 2005).

Wolf-Rayet (W-R) galaxies are natural test beds for the study of stellar chemical evolution models for the enhancement of CNO elements in massive stars. W-R galaxies are a subset of the class of starburst galaxies or emission-line and H II galaxies, whose integrated spectra present the “starprint” of W-R stars,

i.e., the broad emission spectral features associated with W-R stars, the main feature of which is the broad He II $\lambda 4686$ emission line (e.g., Osterbrock & Cohen 1982; Kunth & Joubert 1985; Conti 1991; and López-Sánchez & Esteban 2008 among others). W-R stars are chemically evolved end stages of the most massive stars within a starburst region (Crowther 2007). They have very short lives (about 10^5 yr), hence they can only be detected in population when numerous. This implies that the starburst activity of W-R stars is dominant with respect to the lower-mass stars. Single stars with masses greater than 30 to $60 M_{\odot}$ (depending on the metallicity and rotation rate) become W-R stars (Maeder & Meynet 1994). Therefore, the “starprint” of W-R stars indicates a top-heavy initial mass function (IMF). López-Sánchez & Esteban (2008, 2009, 2010a, 2010b) and López-Sánchez (2010) conducted the hereto most complete observational study of W-R galaxies. Their observations include ground-based optical spectra, deep broad- and narrow-band images, radio, and X-rays. From here on we will refer to the works of López-Sánchez & Esteban (2008, 2009, 2010a) and López-Sánchez & Esteban (2010b) as LSE08, LSE09, LSE10a, LSE10b, respectively.

This study has two main motivations: (i) to determine the source of most of the carbon production (i.e., either massive or intermediate-mass stars), and (ii) to study the behavior of carbon as a function of chemical composition. This information will allow better constraints on the stellar and galactic models of chemical evolution. To address these points, we used low-resolution *Hubble Space Telescope* (HST) Space Telescope Imaging Spectrograph (STIS) long-slit spectra of 18 local starburst galaxies. This work is divided into the following sections: the sample selection is described in Section 2; the observation details, a general description of the sample, and data reduction are described in Section 3; the data analysis and methodology, including line flux analysis and reddening correction, a brief description of the Direct Method, the physical conditions of the ionized gas, and the chemical composition are presented in Section 4; and the discussion, and summary and conclusions are presented in Sections 5 and 6, respectively. The Markov Chain Monte Carlo (MCMC) modeling of photoionized objects is presented in the Appendix.

2. Description of the Sample

The sample for this paper was drawn from the W-R galaxies studied by LSE08. Their original sample included 20 galaxies; however, we removed two objects: NGC 5253, since HST archival data already exist (Kobulnicky et al. 1997), and SBS 1211+540, because its faintness required prohibitively long exposure times. This section describes each of the 18 objects of our W-R galaxy sample. The main properties of each object are presented in Table 1. We follow LSE08, LSE09, LSE10a, LSE10b, and López-Sánchez (2010) for most of the general information of each galaxy. We have verified that the regions of the objects observed in our HST STIS data were the same as those observed in the works of López-Sánchez & Esteban.

2.1. Mrk 960

Mrk 960 was cataloged by Haro (1956) as Haro 15, a blue galaxy with emission lines. This object has been extensively studied in a wide range of wavelengths, i.e., in the UV (Kazarian 1979; Kinney et al. 1993; Heckman et al. 1998), in the optical (Cairós et al. 2001a, 2001b; LSE08; LSE09; LSE10a; LSE10b;

Firpo et al. 2011; Hägele et al. 2012), in the NIR (Coziol et al. 2001; Dors et al. 2013), in the far-IR (Calzetti et al. 1994, 1995), and in radio (Gordon & Gottesman 1981; Klein et al. 1984, 1991). Our STIS observations correspond to the center region, “C,” described in LSE08; in that work, Mrk 960 is referred to as Haro 15.

2.2. SBS 0218+003

SBS 0218+003 is included in the W-R galaxies catalog of Schaerer et al. (1999). It is the most distant object analyzed in this work as well as in LSE08, LSE09, LSE10a, and LSE10b, in which the object is referred to as UM 420. A note provided in the NASA/IPAC Extragalactic Database (NED) indicates that SBS 0218+003 is probably an H II region in UGC 1809. López-Sánchez et al. compared the spectra of SBS 0218+003 with that of UGC 1809 and concluded that the latter is an S0 spiral galaxy at redshift $z = 0.0243$. By comparing the radial velocities of both objects, they concluded that they are not physically related. In the work of LSE08, SBS 0218+003 is referred to as UM 420.

2.3. Mrk 1087

Mrk 1087 was classified by Conti (1991) as an emission-line galaxy without the broad emission line He II 4686 Å, and it was later classified by López-Sánchez et al. (2004b) as a luminous blue compact galaxy (BCG) within a group of interacting dwarf objects. These authors argue that Mrk 1087 does not host an active galactic nucleus, and that this galaxy and its dwarf companions should be considered to be a group of galaxies. According to López-Sánchez et al. (2004b), the various filaments of Mrk 1087 and surrounding dwarf objects suggest that this could be a group in interaction. Such filaments were first reported by Méndez & Esteban (2000). Our STIS observations of Mrk 1087 correspond to the center knot of López-Sánchez et al. (2004b) and LSE08.

2.4. NGC 1741

NGC 1741 is the brightest member of the interacting group of galaxies HCG 31. According to López-Sánchez et al. (2004a), the analysis of the kinematics of HCG 31 suggests that an almost simultaneous interaction involving several objects are taking place. In the nomenclature given by Hickson (1982), NGC 1741 actually corresponds to HCG 31C; nonetheless, since objects A and C are clearly interacting, the two objects can be considered a single entity called HCG 31 AC. A detailed analysis of these interacting galaxies in broad-band imaging and optical intermediate-resolution spectroscopy is presented in López-Sánchez et al. (2004a). Our STIS observations correspond to the northeast part of knot AC in López-Sánchez et al. (2004a) and LSE08.

2.5. Mrk 5

Markarian (1967) included Mrk 5 in his first list of galaxies with UV continua; later, it was classified by Conti (1991) as an emission-line galaxy with a narrow He II 4686 Å in emission. It is usually classified as a H II galaxy and/or a cometary-type Blue Compact Dwarf Galaxy (BCDG). It has an extensive, regular, and elliptical envelope formed by old stars, and it is a low-metallicity object (LSE08). Our STIS observations of Mrk 5 correspond to the slit position INT-1 in the work of LSE08.

2.6. Mrk 1199

Mrk 1199 is part of a group of interacting galaxies. The main body of the group is an Sb galaxy, which is interacting with an elliptical object located to the NE of the main galaxy. Mrk 1199 was classified as a W-R galaxy by Schaerer et al. (1999). The [O III] 4363 Å line was reported as not detected in the works of Izotov & Thuan (1998) and LSE09. We did not detect this line either; however, we did observe a nebular He II 4686 Å emission line. Our STIS observations of Mrk 1199 correspond to slit position D in the work of LSE08.

2.7. IRAS 08208+2816

IRAS 08208+2816 is classified as a H II galaxy, and it is included in the catalog of W-R galaxies of Schaerer et al. (1999). Huang et al. (1999) first reported both nebular and broad He II 4686 Å emission lines, as well as W-R blue and red bumps (at C III 4650 Å and C IV 5808 Å, respectively), which suggest the presence of populations of late-type WN stars (WNL) and early-type WC stars (WCE) in the galaxy. Our STIS observations of IRAS 08208+2816 correspond to knot C in LSE08.

2.8. IRAS 08339+6517

IRAS 08339+6517 is a luminous infrared and Ly α -emitting starburst galaxy, cataloged as a W-R galaxy by López-Sánchez et al. (2008). These authors presented a detailed study of the deep broad-band optical images, narrow-band H α CCD images, and optical intermediate-resolution spectra of IRAS 08339+6517 and its dwarf companion. They concluded that the chemical composition of both galaxies is similar, and that these objects are most likely kinematically interacting. Our STIS observations of IRAS 08339+6517 correspond to knot A in LSE08.

2.9. SBS 0926+606A

SBS 0926+606A is one component of the pair of objects of SBS 0926+606, where component A is a BCDG and component B is a more elongated object north of object A and with no W-R features detected. Izotov et al. (1994) first detected the narrow He II 4686 Å emission line in SBS 0926 + 606A to measure the primordial helium abundance. The galaxy was later studied spectroscopically by several authors (e.g., Izotov et al. 1997; Pérez-Montero & Díaz 2003; Kniazev et al. 2004), and the properties of massive stars in this galaxy were studied by Guseva et al. (2000). Our STIS observations of SBS 0926+606 correspond to knot A in LSE08.

2.10. Arp 252

Arp 252 is classified as an interacting pair of galaxies. Our STIS observations correspond to the brighter galaxy, ESO 566–8, which is the northern object. Most of the H α emission of the entire system (93%) comes from the brighter galaxy (LSE08). Our STIS observations of Arp 252 correspond to knot A (ESO 566–8) in LSE08.

2.11. SBS 0948+532

SBS 0948+532 is an emission-line galaxy and a BCDG included in the W-R galaxies catalog of Schaerer et al. (1999).

Table 1
Main Properties of the W-R Galaxy Sample Analyzed in This Study

Galaxy Name	R.A. (2000)	Decl. (2000)	$E_{(B-V)G}^a$	d (Mpc) ^b	Other Names
Mrk 960	00 48 35.4	-12 42 60	0.021	85.9 ± 6.0	Haro 15, LEDA 2845, IRAS 00460-1259
SBS 0218+003	02 20 54.5	+00 33 23	0.032	235.6 ± 16.5	UM 420
Mrk 1087	04 49 44.5	+03 20 03	0.056	112.1 ± 7.8	II Zw 23, IRAS 04470+0314
NGC 1741	05 01 37.4	-04 15 31	0.045	54.6 ± 3.8	HCG 31 AC, Mrk 1089, SBS 0459-043
Mrk 5	06 42 15.5	+75 37 30	0.076	13.2^c	SBS 0720+756, LEDA 19459
Mrk 1199	07 23 43.5	+33 26 31	0.047	57.7 ± 4.0	SBS 0720+20335, LEDA 20911
IRAS 08208+2816	08 23 54.9	+28 06 22	0.028	193.9 ± 13.6	WR 238
IRAS 08339+6517	08 38 23.1	+65 07 15	0.083	81.5 ± 8	LEDA 24283
SBS 0926+606A	09 30 06.4	+60 26 53	0.027	60.2 ± 4.2	LEDA 26955
Arp 252	09 44 59.7	-19 42 46	0.043	129.8^d	ESO 566-7 + ESO 566-8
SBS 0948+532	09 51 32.0	+52 59 36	0.011	192.9 ± 14	LEDA 28398
Tol 9	10 34 38.7	-28 35 00	0.058	45.5 ± 3.2	Tol 1032-283, ESO 435-42
SBS 1054+365	10 57 47.0	+36 15 39	0.018	10.5^c	LEDA 32954
POX 4	11 51 11.6	-20 35 57	0.035	52.5 ± 3.9	Tol 1148-2020, LEDA 37074
SBS 1319+579	13 21 22.6	+57 41 29	0.012	35.1 ± 2.5	NGC 5113
SBS 1415+437	14 17 01.4	+43 30 5	0.008	10.1^c	LEDA 51017, CG 413
Tol 1457-262	15 00 29.0	-26 26 49	0.138	81.3 ± 5.9	ESO 513-11, IRAS 1457-2615
III Zw 107	23 30 09.8	+25 31 50	0.053	77.9 ± 5.5	IV Zw 153, LEDA 71605, IRAS 23276+2515

Notes.

^a Galactic reddening values calculated from the NASA/IPAC Extragalactic Database (NED) with $R_V = 3.1$.

^b Distance taken from NED with $H_0 = 73 \pm 5 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and the Virgo-GA-Shapley model.

^c Distances taken from Zhao et al. (2013).

^d Distance taken from LSE08.

Izotov et al. (1994) first detected the He II 4686 Å emission line in this object. Guseva et al. (2000) reanalyzed SBS 0948+532 and found the presence of WNL stars and the tentative evidence of a red W-R bump.

2.12. Tol 9

Tol 9 is the most metal-rich object in our sample. It is classified as an emission-line galaxy without the emission line He II 4686 Å, and as a W-R galaxy by Schaerer et al. (1999). Wamsteker et al. (1985) suggested that Tol 9 is interacting with a nearby object. Our STIS observations of Tol 9 correspond to slit position INT in the work of LSE08.

2.13. SBS 1054+365

SBS 1054+365 is a BCDG included in the W-R galaxies catalog of Schaerer et al. (1999) and in the catalog of interacting galaxies of Vorontsov-Velyaminov (1959, 1977), due to the detection of a nearby companion about 1 arcminute to the north. Our STIS observations focused on region C, which is the brightest knot; these observations correspond to the main component in the work of LSE08.

2.14. POX 4

POX 4 is included in the catalog of W-R galaxies of Conti (1991) as well as in the catalog of Schaerer et al. (1999). The broad He II 4686 Å emission line was first detected in POX 4 by Kunth & Joubert (1985). It is classified as a BCDG with its bright knot surrounded by three or four star-forming regions. López-Sánchez & Esteban (2010a) detected both broad He II 4686 Å and C IV 5808 Å to determine the number of WNL and WCE stars. Our STIS observations of POX 4 correspond to the main component in the work of LSE08.

2.15. SBS 1319+579

The He II 4686 Å emission line was first detected in the BCDG SBS 1319+579 by Izotov et al. (1994). Schaerer et al. (1999) included this object in their W-R galaxies catalog, and later Guseva et al. (2000) detected in it WNL and WCE populations. Our STIS observations correspond to knot A in LSE08, which is the brightest knot in SBS 1319+579.

2.16. SBS 1415+437

SBS 1415+437 is the most metal-poor object in our sample, and it is one of the most metal-poor BCDGs known. From broad-band photometry, LSE08 found that the brightest regions of the galaxy are blue, yet they are slightly redder when considering the flux from the entire galaxy, suggesting the existence of a low-luminosity component dominated by older stellar populations. Our STIS observations of SBS 1415+437 correspond to knot A in the work of LSE08.

2.17. Tol 1457-262

Tol 1457-262 is classified as a pair of galaxies with significant star formation activity. W-R features have been detected in the western object by several authors (e.g., Conti 1996; Pindao 1999; LSE08; LSE09; LSE10a; LSE10b; Esteban et al. 2014). Our STIS observations correspond to region B in LSE08. Tol 1457-262 is included in the W-R galaxies catalog of Schaerer et al. (1999). Our STIS observations of Tol 1457-262 correspond to knot A in the work of LSE08.

2.18. III Zw 107

III Zw 107 is classified as an emission-line galaxy. It was named after the Catalog of Selected Compact Galaxies and of Post-Eruptive Galaxies by Zwicky (1971). Photometric studies

Table 2
HST STIS Observation Log of the W-R Galaxy Sample Analyzed in This Study

Galaxy Name	Observation Date	Position Angle ^a (deg)	Orbits	Exposure Time [s] per Setting			
				Gratings		Mirror ^b G750L	
				G230L	G430L		
Mrk 960	2012 Sep	−76	1	1270	300	150	20
SBS 0218+003	2012 Dec	29	2	3915	500	250	20
Mrk 1087	2012 Jan	−3	1	1315	300	150	20
NGC 1741	2012 Jan	10	1	1155	300	150	60
Mrk 5	2012 Nov	−180	1	1555	300	150	60
Mrk 1199	2012 Nov	−144	1	1420	300	150	5
IRAS 08208+2816	2012 Jan	−162	1	1300	300	150	30
IRAS 08339+6517	2013 Jan	144	1	1720	300	150	10
SBS 0926+606A	2013 Feb	117	2	4250	600	300	60
Arp 252	2013 Mar	18	1	1370	300	150	10
SBS 0948+532	2013 Mar	117	4	8550	1842	900	120
Tol 9	2012 May	49	2	4085	500	250	10
SBS 1054+365	2013 May	59	2	4090	500	250	30
POX 4	2012 Apr	25	1	1170	300	150	60
SBS 1319+579	2012 Aug	39	4	8745	1899	912	120
SBS 1415+437	2014 Jan	−133	2	4170	500	250	30
Tol 1457−262	2012 Jun	15	1	1340	300	150	20
III Zw 107	2012 Jan	0	1	1340	300	150	20

Notes.

^a Position angles are with reference to the north (upper edge of acquisition images, with east—positive—toward the left side).

^b Seconds to obtain the acquisition images.

have been performed on III Zw 107 by Moles et al. (1987) and Cairós et al. (2001a, 2001b). Spectrophotometric studies of this galaxy in the visual, X-ray, and radio have been performed by LSE08, LSE09, LSE10a, and LSE10b. Kunth & Joubert (1985) included III Zw 107 in their catalog of W-R galaxies. Our STIS observations of III Zw 107 correspond to knot A in the work of LSE08.

3. Observations

The sample was observed in *HST* program GO 12472 (PI: Leitherer), which uses STIS to perform co-spatial spectroscopy over the wavelength range of 1600–10000 Å. The final coordinates are given in Table 2, after target acquisition of the telescope. Distances higher than 20 Mpc were taken from NED with the Hubble flow calculations assuming $H_0 = 73 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and the Virgo–GA–Shapley; closer distances were taken from the work of Zhao et al. (2013).

The observation program was conducted between January 2012 and January 2014. We used the long-slit NUV-MAMA and CCD detectors with three gratings: G230L for the MAMA detector, and G430L and G750L for the CCD. The G230L grating has a spectral range from 1560 to 3180 Å and an average dispersion of $1.58 \text{ Å pixel}^{-1}$; the G430L grating has a spectral range from 2900 to 5700 Å and an average dispersion of $2.73 \text{ Å pixel}^{-1}$; and the G750L grating has a spectral range from 5240 to 10270 Å and an average dispersion of $4.92 \text{ Å pixel}^{-1}$. All three gratings have a resolving power, R , of about 500. The properties of the observations are described in Table 2. The spectra were taken with the $0''.2 \times 52''$ aperture, which is a good compromise between slit loss and spectral resolution at $R \sim 500$. Prior to taking the STIS spectra, we used the CCD detector to obtain a 5×5 arcsecond (or 100×100 pixels) target acquisition image. We used this image to check the acquisition of the STIS spectra. Figures 1 and 2 show the acquisition

images. The slit positions are just as taken from the proposal. The *HST* data used for this analysis can be downloaded from the Mikulski Archive for Space Telescopes (MAST; <https://doi.org/10.17909/T96S3J>).

3.1. Data Reduction

The data were processed with the CALSTIS pipeline (Biretta et al. 2015), which includes the following steps: conversion from high-resolution to low-resolution pixels (MAMA), linearity correction, dark subtraction, cosmic-ray rejection (CCD), combination of cr-split images (CCD), flat-fielding, geometric distortion correction, wavelength calibration, and photometric calibration. We extracted one-dimensional (1D) spectra of each object from the x2d (MAMA) and sx2 (CCD) files, which contain two-dimensional (2D) spectral images. Since some of our objects are very faint, we re-extracted all spectra with four different extraction windows: 11, 16, 21, and 30 pixels. The default pipeline for the extended objects is 11 pixels for the MAMA spectra and 7 pixels for the CCD. For each NUV, optical, and NIR spectra, we chose the extraction window with the best signal-to-noise ratio (S/N). The CCD spectra were also cleaned of cosmic rays using the Python module `cosmics` based on Pieter van Dokkum’s L.A.Cosmic (van Dokkum 2001). We show the spectra of our best and worst S/N objects, POX 4 (Figure 3) and Mrk 1199 (Figure 4), respectively.

4. Data Analysis and Methodology

We followed traditional analysis using a two-zone approximation to define the temperature structure of the object and used it to calculate the oxygen abundances. We obtained the C/H abundances with the method described in Garnett et al. (1995). To determine if our carbon and oxygen abundances were sensible, we used an MCMC method to probe the

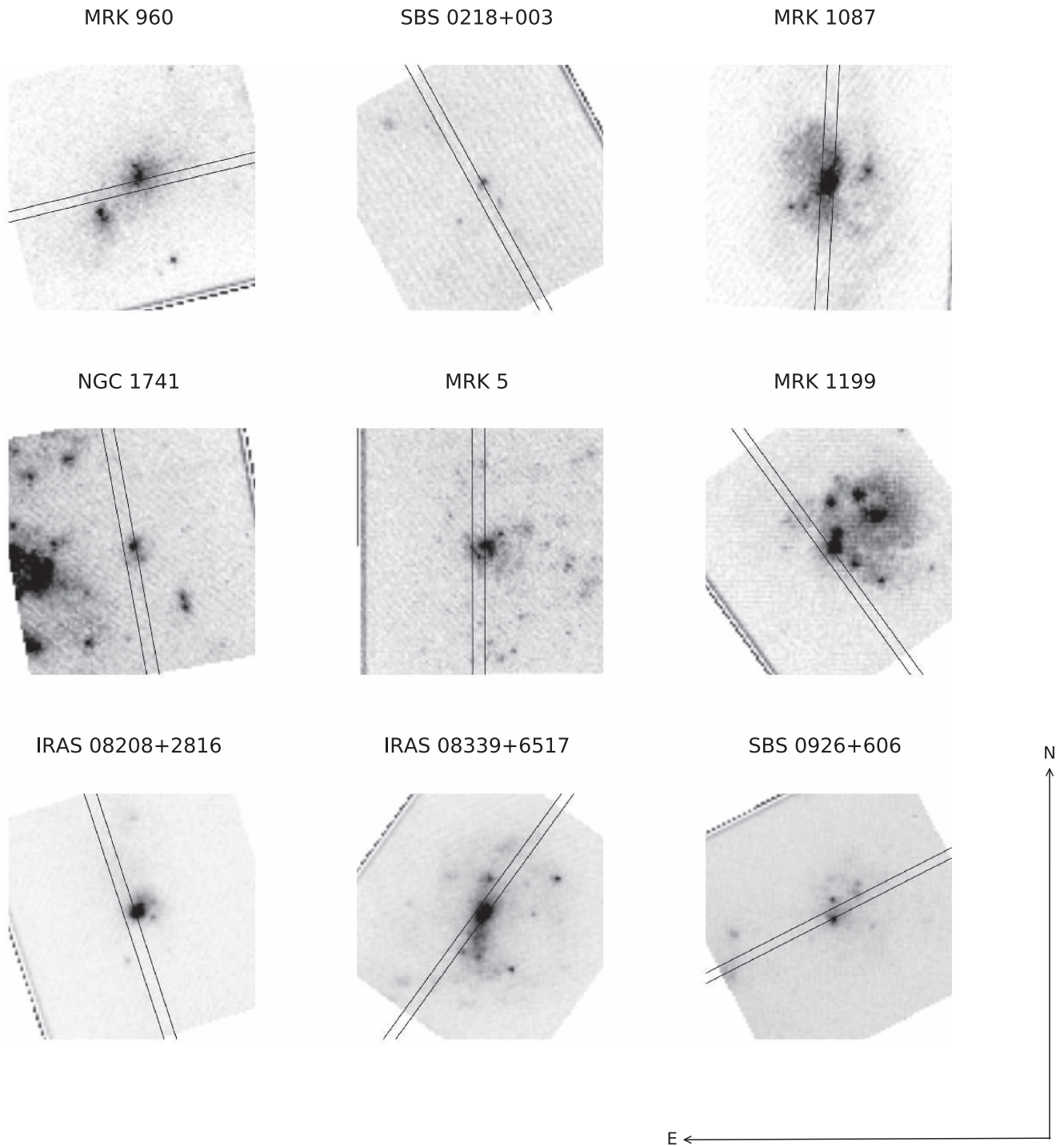


Figure 1. Optical *HST* target acquisition images rotated north and showing the slit positions for our STIS observations for the first half of the sample. Each image is 5×5 arcseconds or 100×100 pixels, hence the 0.2×52 slit used corresponds to 4 pixels. The angles presented in Table 2 are with reference to the north, indicated in the lower-right side of the image; positive angles are to the east. The origin of the coordinate system is at the center of each acquisition image.

parameter space. This methodology is described in detail in the [Appendix](#).

4.1. Line Flux Analysis and Reddening Correction

Line flux measurement was done with a Python code we developed. This code determines the stellar continuum, finds the emission lines from a catalog we compiled, and measures the total flux. This catalog contains typical emission and absorption lines observed in nebular spectra, composed as

follows: lines from about 1150 to 2850 Å were taken from Leitherer et al. (2011), and lines from about 3200 to 10300 Å were taken from Peimbert (2003). We measured all lines with width at the continuum greater than 1.5 \AA . The continuum was determined by sigma-clipping the strong emission lines and then finding the flux mode of the remainder signal. The flux of the emission lines was determined with a simple sum of the flux over continuum routine.

To obtain the reddening corrected-intensities, we included PyNeb v.0.9.13 (Luridiana et al. 2015) in our code. We used

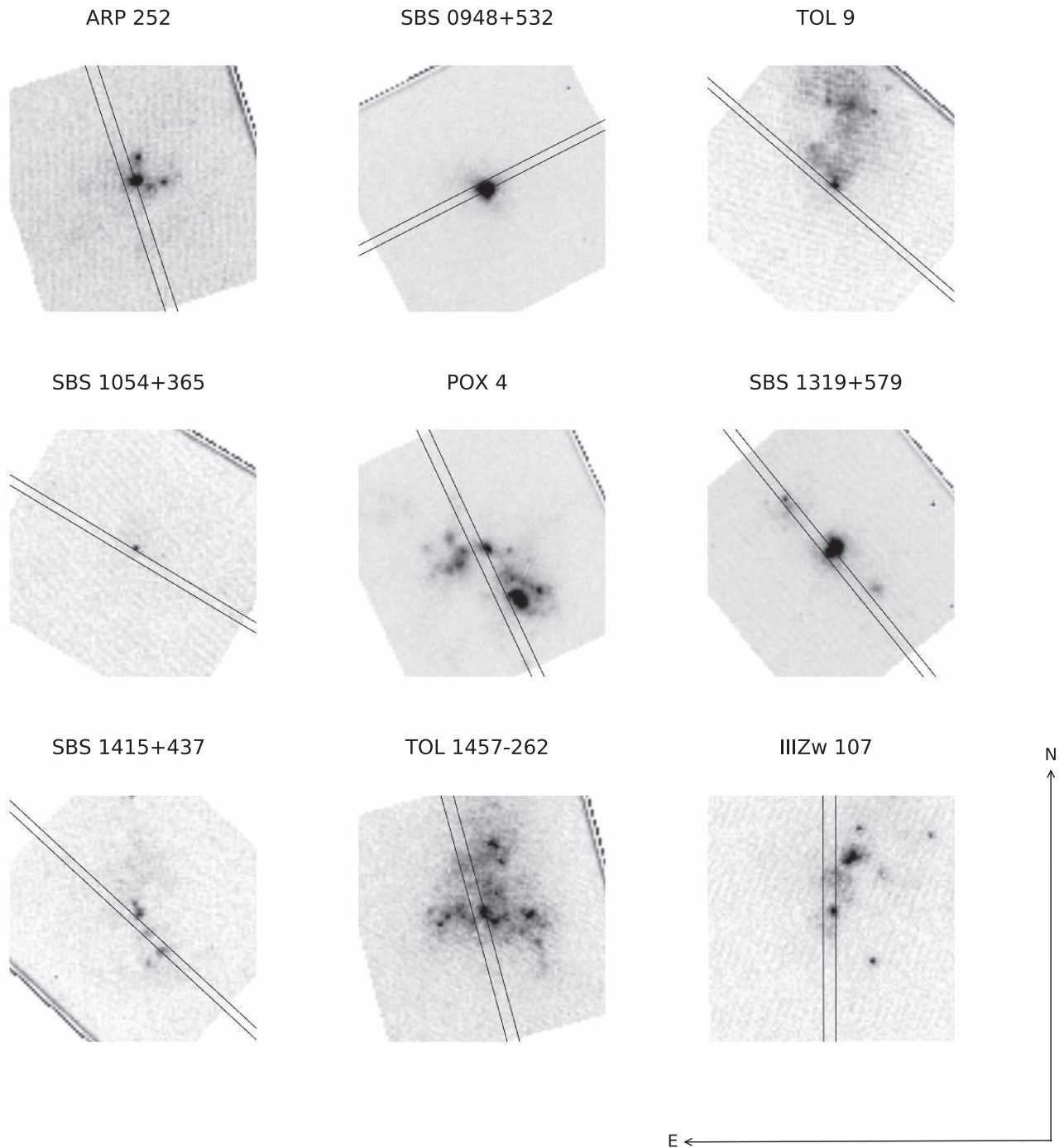


Figure 2. Optical *HST* target acquisition images rotated north and showing the slit positions for our STIS observations for the second half of the sample. Image and slit sizes are the same as in Figure 1.

the extinction law by Fitzpatrick (1999) for the UV, and that by Fitzpatrick & Massa (1990) for the optical and NIR, both with $R_V = 3.1$. The Balmer and helium emission lines were corrected for underlying absorption, and the adopted EWs in absorption were taken from a stellar spectra template normalized to EW_{abs} provided in Table 2 of Peña-Guerrero et al. (2012b). This template was based on the low-metallicity instantaneous bursts models from González-Delgado et al. (1999), as well as additional models ran by M. Cerviño using the same code as González-Delgado and collaborators. In order to find the values of $C(\text{H}\beta)$ for this sample, since $\text{H}\alpha$ is blended with the $[\text{N II}]$ lines due to its low resolution ($R \sim 500$), we

performed an iterative process using the expected theoretical value of $\text{H}\alpha$ according to Storey & Hummer (1995); the measured fluxes of $\text{H}\beta$, $\text{H}\gamma$, and $\text{H}\delta$; and the deblend task in the Pyraf routine `splot`. The deblended intensity of $\text{H}\alpha$ is presented in Tables 3–5.

The dereddened line intensities and final values used for $C(\text{H}\beta)$ and EW_{abs} , as well as the equivalent widths for the other important lines, are presented in Tables 3–5. The structure is the same for all three tables: columns 1 and 2 are, respectively, the rest-frame wavelength and the line identification (ID); column 3 is the reddening law used (f_λ values), and columns 4 through 9 show the dereddened line intensities

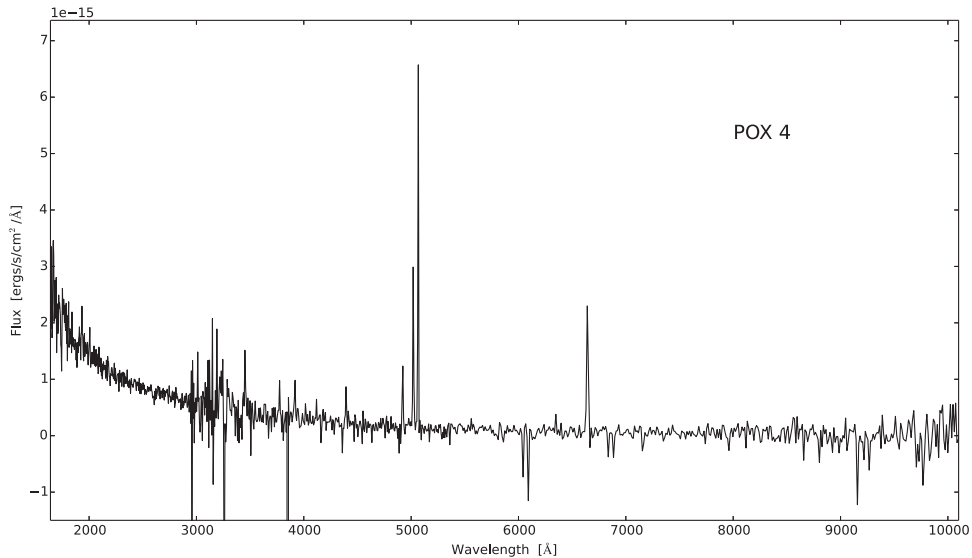


Figure 3. Combined NUV, optical, and NIR *HST* STIS spectra of POX 4.

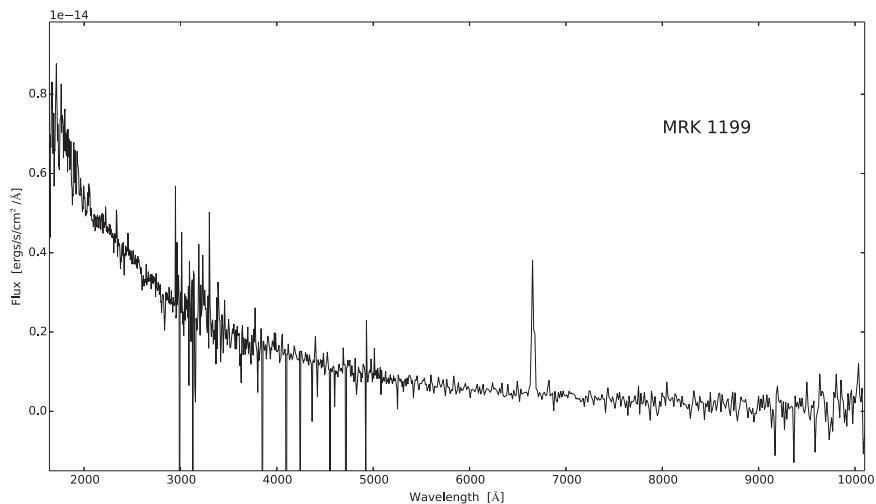


Figure 4. Combined NUV, optical, and NIR *HST* STIS spectra of Mrk 1199.

relative to $H\beta$, with the standard assumption that $I(H\beta) = 100$. The values of $C(H\beta)$ that we obtained agree within the errors with the values determined by LSE09.

To obtain the uncertainties of the measured lines, we used the estimated nominal spectroscopic accuracies for flux calibration in L mode given in the STIS Data Handbook: 2%, 5%, and 5% for NUV, optical, and NIR, respectively. We assumed the uncertainties to be symmetric around the central wavelength. The contribution to the uncertainties due to the noise was estimated from the rms of the continuum adjacent to the emission line. The final adopted uncertainties were estimated using standard error propagation equations.

4.2. Direct Method

The so-called direct method assumes a homogeneous temperature structure throughout the whole volume of the object, and then this temperature is used to determine the abundances of all available ions. This method generally adopts a two-ionization zone approximation, where the temperature of the high-ionization zone can be represented by the electron

temperature of [O III], T_e [O III], or T_e [S III], and the temperature of the low-ionization zone can be represented by T_e [O II] or T_e [N II].

4.3. Physical Conditions of the Ionized Gas

To corroborate the fact that photoionization in our W-R galaxy sample is caused by massive stars, we created an [O III]/ $H\beta$ to [N II]/ $H\alpha$ diagram, commonly referred to as the BPT (Baldwin et al. 1981) diagram. We were not able to separate the [N II] lines from $H\alpha$; nonetheless, we used the [N II] line intensities presented in López-Sánchez et al. (2004a, 2004b, 2008) and LSE09. It is important to note that our observations seem to correspond to the center regions observed in the works of LSE, however the exact location of the slits most likely changed. Figure 5 shows the [O III]/ $H\beta$ to [N II]/ $H\alpha$ observed values for our sample, as well as the intensity ratios measured by López-Sánchez et al. We also show the theoretical upper limit for starburst galaxies as given in Kewley et al. (2001) as a dotted green line, the lower limit for active galactic nuclei (AGN) as presented in Kauffmann et al. (2003) as a dashed-dotted

Table 3
Line Intensity Ratios for Mrk 960, SBS 0218+003, Mrk 1087, NGC 1741, Mrk 5, and Mrk 1199

λ (Å)	Id.	f_λ	$I(\lambda)/I(H\beta), I(H\gamma) = 100$					
			Mrk 960	SBS 0218+003	Mrk 1087	NGC 1741	Mrk 5	Mrk 1199
1661	O III]	1.124	...	4 ± 2	...	1:
1666	O III]	1.122	13 ± 6	...	1:	...	4 ± 2	1:
1883	Si III]	1.187	12 ± 1	6:	29 ± 3	9 ± 1	9 ± 1	8:
1892	Si III]	1.197	16 ± 2	1:	15 ± 1	...
1902	Fe II	1.208	14 ± 1	...	22 ± 2	8 ± 1	17 ± 1	13 ± 4
1909	C III]	1.214	85 ± 32	8 ± 2	5 ± 2	11 ± 2	10 ± 4	5 ± 1
2322	[O III]	1.337	...	6:	3:	3:	5:	3:
2471	[O II]	0.999	1:	2:	1:	6 ± 1	1:	6:
2569	[Na VI]	0.858	8 ± 1	...	7 ± 1	3:	1:	...
2661	[Al II]	0.764	6 ± 1	2:	...	2:	1:	5:
2670	[Al II]	0.757	...	1:	1:	2:	...	5:
2783	[Mg V]	0.670	1:	6:	9 ± 1	5:	...	4:
2854	[Ar IV]	0.623	1:	...	3:	1:
2868	[Ar IV]	0.615	2:	6:	...	5:	4:	7 ± 2
2871	[Na VI]	0.613	4:	7:	...	6 ± 1	3:	7 ± 2
2928	[Mg V]	0.580	7 ± 1	10 ± 1	12 ± 1	4:	8:	4 ± 2
2970	[Na VI]	0.558	5:	3:	9 ± 1	3:
3242	[Na IV]	0.440	56 ± 6	18 ± 2	...	136 ± 12	99 ± 7	4:
3362	[Na IV]	0.398	54 ± 6	1:	107 ± 11	107 ± 9	...	9:
3727	[O II]	0.291	99 ± 9	188 ± 14	149 ± 14	268 ± 23	111 ± 8	93 ± 18
3798	H 10+[S III]	0.272	35 ± 4	6 ± 1
3869	[Ne III]	0.254	21 ± 3	26 ± 2	18 ± 6	15 ± 2	56 ± 4	4:
3889	H 8+He I	0.248	36 ± 4	27 ± 2	17 ± 2	14 ± 2	15 ± 1	8:
3967	[Ne III]	0.229	22 ± 3	15 ± 1	4 ± 1	12 ± 1	...	25 ± 4
4069	[S II]	0.203	22 ± 3	3:	...	14 ± 2	...	18 ± 4
4076	[S II]	0.202	31 ± 4	2:	...	11 ± 2	...	8 ± 3
4102	Hδ	0.195	26 ± 3	15 ± 1	25 ± 3	25 ± 2	14 ± 2	23 ± 4
4125	[Fe II]	0.190	18 ± 2	8 ± 1	...	7 ± 1	...	10 ± 4
4340	Hγ	0.137	48 ± 5	48 ± 4	38 ± 5	46 ± 4	47 ± 3	47 ± 8
4363	[O III]	0.132	1:	4:	35:	...	5 ± 1	<5
4563	Mg I]	0.081	7 ± 2	8 ± 1	4:
4571	Mg I]	0.079	21 ± 3	4:	2:	2 ± 1	...	4:
4658	[Fe III]	0.056	12 ± 2	4:	36 ± 5	19 ± 2	19 ± 1	8 ± 3
4686	He II	0.048	...	4:	...	<4	...	<5
4702	[Fe III]	0.044	35 ± 4	3:	14 ± 3	6 ± 1	...	4:
4711	[Ar IV]+He I	0.041	26 ± 4	...	22 ± 2	4:
4740	[Ar IV]	0.034	5 ± 2	...	11 ± 3	3 ± 1	...	5:
4755	[Fe III]	0.029	8 ± 2	1:	4:	41 ± 4	...	4:
4861	Hβ	0.000	100 ± 12	100 ± 8	100 ± 11	100 ± 9	100 ± 7	100 ± 14
4922	He I	-0.017	23 ± 3	3:	22 ± 4	2 ± 1	...	6 ± 1
4959	[O III]	-0.027	67 ± 8	110 ± 7	29 ± 4	60 ± 6	128 ± 11	34 ± 8
5007	[O III]	-0.040	217 ± 23	362 ± 27	98 ± 13	172 ± 15	391 ± 28	65 ± 12
5041	Si III]	-0.049	3 ± 1	8 ± 1	8 ± 2	...	25 ± 2	1:
5056	Si III]	-0.053	9 ± 2	23 ± 2	7 ± 2	...	26 ± 2	4:
5192	[Ar III]	-0.089	6 ± 1	13 ± 1	8 ± 3	14 ± 1	...	10 ± 2
5198	[N I]	-0.091	3:	26 ± 2	17 ± 3	5 ± 1	8 ± 1	4:
5518	[Cl III]	-0.165	26 ± 3	18 ± 2	29 ± 4	...	8 ± 1	9 ± 2
5538	[Cl III]	-0.170	37 ± 4	18 ± 1	65 ± 7
5755	[N II]	-0.213	35 ± 4	...	4:	13 ± 1	8 ± 1	2:
5876	He I	-0.235	14 ± 3	11 ± 1	14 ± 2	51 ± 5	10 ± 1	7 ± 3
6312	[S III]	-0.309	2:	8 ± 1	2:	...
6563	Hα	-0.348	249 ± 27	287 ± 23	286 ± 31	286 ± 26	286 ± 21	287 ± 39
6678	He I	-0.365	5:	27 ± 3	...	6 ± 2
6716	[S II]	-0.371	36:	28 ± 2	59 ± 7	24 ± 2	14 ± 1	16 ± 4
6731	[S II]	-0.373	54 ± 13	30 ± 2	46 ± 5	18 ± 1	17 ± 1	18 ± 2
7065	He I	-0.419	6 ± 1	5:	...	2:	...	3 ± 1
7136	[Ar III]	-0.428	21 ± 2	13 ± 1	12 ± 2	18 ± 2	6:	2 ± 1
7281	He I	-0.446	8 ± 2	17 ± 1	9 ± 2
7320	[O II]	-0.451	20 ± 4	28 ± 2	166 ± 19	33 ± 3	...	19 ± 3
7330	[O II]	-0.452	3:	21 ± 2	82 ± 9	24 ± 2	...	11 ± 2
7751	[Ar III]	-0.500	5 ± 1	41 ± 3	4 ± 1
8481	[Cl III]	-0.571	36 ± 4	...	38 ± 4	33 ± 3	27 ± 2	6 ± 1
9069	[S III]	-0.617	35 ± 3	64 ± 4	23 ± 2	19 ± 2

Table 3
(Continued)

λ (Å)	Id.	f_λ	$I(\lambda)/I(\text{H}\beta), I(\text{H}\beta) = 100$					
			Mrk 960	SBS 0218+003	Mrk 1087	NGC 1741	Mrk 5	Mrk 1199
9531	[S III]	-0.649	58 ± 6	22 ± 3	374 ± 33	36 ± 3	41 ± 3	22 ± 3
$F(\text{H}\beta)^a$			2.38 ± 0.20	3.58 ± 0.19	2.21 ± 0.17	5.76 ± 0.36	4.96 ± 0.26	1.82 ± 0.19
$C(\text{H}\beta)$			0.03 ± 0.02	0.38 ± 0.05	0.14 ± 0.05	0.01 ± 0.05	0.19 ± 0.05	0.23 ± 0.05
EW_{abs}			0.5 ± 0.3	2.0 ± 0.4	2.0 ± 0.4	2.0 ± 0.3	1.0 ± 0.4	0.3 ± 0.3
$EW(\text{H}\alpha)^b$			97 ± 6	761 ± 26	57 ± 3	432 ± 24	1458 ± 73	148 ± 13
$EW(\text{H}\beta)^b$			22 ± 2	269 ± 12	15 ± 2	63 ± 3	404 ± 17	20 ± 2
$EW([\text{O III}] 5007)^b$			53 ± 3	1069 ± 47	18 ± 2	123 ± 5	1619 ± 68	16 ± 2
$EW(\text{C III}] 1909)^b$			1:	2:	1:	2:	3 ± 1	1:

Notes.

^a Units of $F(\text{H}\beta)$ are given in $\text{erg s}^{-1} \text{cm}^{-2} \times 10^{-15}$.

^b Units of equivalent widths (EWs) are given in Å.

magenta line, and the division between AGN and low-ionization nuclear emission-line regions (LINERs) according to Kauffmann et al. (2003) as a dashed magenta line. The bulk of the objects from our W-R galaxy sample fall on the star formation region of the diagram, though there are two objects (SBS 1319+579 and IRAS 08208+2816) that lie in the mix region between the star-forming galaxies and AGNs (Richardson et al. 2016). Mrk 1199 lies a bit farther from the H II region loci than the rest of the sample. This is due to its low-ionization degree (see Section 4.4.3 and values of oxygen ionization degree, OID, and excitation index, P , in Table 10); Sánchez et al. (2015) present a more detailed study on this issue.

We used the direct method to determine temperatures for our sample. For objects where [O III] 4363 Å or [S III] 6312 Å were not observed or did not have good enough signal-to-noise (S/N), we used the temperature presented by LSE09 or LSE10b. We carefully checked that the regions of the objects observed in our *HST* STIS data were the same as those observed in the works of López-Sánchez & Esteban; moreover, we find that our high-ionization zone temperature determinations agree, within the errors, with those obtained by LSE08, LSE09, LSE10a, LSE10b. We were able to obtain a high-ionization zone temperature measurement for 15 out of the 18 objects in our sample. Figure 6 shows our high-ionization zone temperatures versus those obtained by LSE.

The CELs of [O II] 7320 and 7330 Å and/or [S II] 4069 and 4076 Å were not observed or did not have good enough S/N (i.e., intensity uncertainty greater than 70%) to determine temperatures for the low-ionization zone (we defined a line with poor or low S/N as a line with an error greater than 40%). To determine T_e [O II], we used the following relation taken from Garnett (1992):

$$T_e[\text{O II}] = 0.7 \times T_e[\text{O III}] + 3000. \quad (1)$$

Garnett used the linear approximation to the relation between $T(\text{O}^+)$ and $T(\text{O}^{+2})$ provided by Campbell et al. (1986) from the models of Stasińska (1982) to determine the ion-weighted mean electron temperatures.

It is important to keep in mind that the two-zone approximation is indeed a first-order approximation to the actual thermal structure of the nebula (or H II region). For the objects where [O III] 4363 Å or [S III] 6312 Å did not have good enough S/N (i.e., the line had a width greater than 1.5 Å and intensity uncertainty smaller than 50%), we used both $T_e[\text{O III}]$

and $T_e[\text{O II}]$ as presented in LSE10b. The uncertainties in the temperatures and densities were obtained from PyNeb.

To obtain electron density measurements, we used the 6731/6711 [S II] lines. In those objects where we could not determine one of the lines, we adopted the electron density given in LSE09 or LSE10b. We were able to obtain at least an upper electron density limit for 12 out of the 18 objects in our sample. The adopted electron temperatures and densities are presented in Table 6.

4.4. Chemical Abundances

In this section, we describe how we obtained both the ionic and total gas abundances. We also briefly explain corrections to the direct method based on temperature inhomogeneities, dust depletion, and ionization structure.

4.4.1. Ionic Abundances

The ionic chemical abundances of He^+/H^+ , O^{++}/H^+ , O^+/H^+ , $\text{Ne}^{++}/\text{H}^+$, S^{++}/H^+ , and S^+/H^+ were determined with the temperatures and densities shown in Table 6. The resulting ionic abundances are given in Table 7. For the specific case of C^{++}/H^+ , we followed the procedure described in Garnett et al. (1995).

Garnett et al. (1995) used *HST* spectroscopy of dwarf galaxies to measure the relative abundances of $\text{C}^{+2}/\text{O}^{+2}$ from the rest-UV emission lines C III] 1909 Å to O III] 1666 Å and $T_e[\text{O III}]$. The method assumes that the electron density of the H II region in question is well below the critical densities for the collisional de-excitation of both C III] and O III], $n_{\text{crit}} \sim 10^5 \text{cm}^{-3}$ and $n_{\text{crit}} \sim 10^3 \text{cm}^{-3}$, respectively (Osterbrock & Ferland 2005). This is the case for all galaxies in our sample. The abundance of $\text{C}^{+2}/\text{O}^{+2}$ can then be computed in the low-density limit. The total abundance of carbon from the Garnett et al. (1995) method depends on the temperature measured for the high-ionization zone (see Equation (2)), hence the choice of the correct temperature is paramount for an accurate estimation of the total carbon abundance through this method.

The method described in Garnett et al. (1995) is relatively straightforward once the collision strengths for C^{+2} and O^{+2} have been selected; we maintain the values adopted by Garnett et al. (1995). The method then essentially consists of four steps: (i) determine the ionic abundance ratio $\text{C}^{+2}/\text{O}^{+2}$ from the emission-line ratio C III] 1909 Å to O III] 1666 Å, and $T_e[\text{O III}]$,

Table 4
Line Intensity Ratios for IRAS 08208+2816, IRAS 08339+6517, SBS 0926+606A, Arp 252, SBS 0948+532, and Tol 9

λ (Å)	Id.	f_{λ}	$I(\lambda)/I(H\beta), I(H\beta) = 100$					
			IRAS 08208+2816	IRAS 08339+6517	SBS 0926+606A	Arp 252	SBS 0948+532	Tol 9
1661	O III]	1.124	9 ± 3	...
1666	O III]	1.122	1:	1:	8 ± 3	1:	6 ± 1	0.16:
1883	Si III]	1.187	6 ± 1	7 ± 3	7 ± 1	2:	5:	4:
1892	Si III]	1.197	1:	...	2:	4:	9 ± 2	...
1902	Fe II	1.208	5 ± 1	1:	20 ± 2	6:	25 ± 4	8 ± 2
1909	C III]	1.216	6 ± 2	9 ± 3	32 ± 10	5 ± 1	28 ± 3	7 ± 2
2322	[O III]	1.337	2 ± 1	...	4 ± 1	2:	8 ± 2	2:
2471	[O II]	0.999	3:	...	1:	3:	2:	18 ± 4
2569	[Na VI]	0.858	...	2:	...	1:	3:	2:
2661	[Al II]	0.764	4:	6 ± 2	2:	4 ± 1
2670	[Al II]	0.757	...	5 ± 2	2:	2:	2:	11 ± 2
2783	[Mg V]	0.670	3:	...	2:	...	4 ± 1	4 ± 1
2854	[Ar IV]	0.623	4:	1:	...	3 ± 1
2868	[Ar IV]	0.615	4:	2:	...	1:	4 ± 1	6 ± 1
2871	[Na VI]	0.613	4:	1:	...	1:	130 ± 46	6 ± 1
2928	[Mg V]	0.580	5:	4 ± 1	1:	3:	9:	5 ± 1
2970	[Na VI]	0.558	...	5 ± 1	2:	...	6 ± 1	91 ± 20
3242	[Na IV]	0.440	28 ± 3	37 ± 9	179 ± 16	...	5 ± 1	...
3362	[Na IV]	0.398	8 ± 2	43 ± 11	7 ± 1	23:
3727	[O II]	0.291	90 ± 9	169 ± 44	96 ± 9	89 ± 5	100 ± 22	210 ± 61
3798	H 10+[S III]	0.272	10 ± 2	...	28 ± 3	8:	...	19 ± 6
3869	[Ne III]	0.254	47 ± 4	44 ± 12	16 ± 2	20 ± 1	58 ± 13	17 ± 4
3889	H 8+He I	0.248	4 ± 1	6 ± 2	10 ± 1	...	16 ± 3	...
3967	[Ne III]	0.229	15 ± 1	...	28 ± 3	7:	32 ± 7	...
4069	[S II]	0.203	3 ± 1	28 ± 6	22 ± 2	3:	3:	47 ± 13
4076	[S II]	0.202	4 ± 1	...	10 ± 1	2:	2:	13 ± 3
4102	H δ	0.195	26 ± 2	26 ± 5	11 ± 1	...	26 ± 6	27 ± 7
4125	[Fe II]	0.190	1:	44 ± 10	8 ± 1	40 ± 11
4340	H γ	0.137	...	6:	48 ± 5	28 ± 2	47 ± 10	...
4363	[O III]	0.132	3 ± 1	1:	9:	...	9 ± 2	52 ± 16
4563	Mg I]	0.081	...	8 ± 2	4 ± 1	7 ± 2
4571	Mg I]	0.079	6 ± 1	13 ± 2	...	4:
4658	[Fe III]	0.056	...	9 ± 2	27 ± 2	8 ± 1	4 ± 1	5 ± 1
4686	He II	0.048	1:	1:	2:	...
4702	[Fe III]	0.044	3 ± 1	...	8 ± 1	8 ± 1	1:	2:
4711	[Ar IV]+He I	0.041	4 ± 1	17 ± 3	8 ± 1	7 ± 1	1:	...
4740	[Ar IV]	0.034	13 ± 1	...	2:	...
4755	[Fe III]	0.029	6 ± 1	42 ± 10	14 ± 1	8 ± 1	3:	16 ± 5
4861	H β	0.000	100 ± 8	100 ± 20	100 ± 9	100 ± 7	100 ± 22	100 ± 21
4922	He I	-0.017	...	7 ± 2	...	5:	...	25 ± 6
4959	[O III]	-0.027	197 ± 17	65 ± 21	191 ± 15	32 ± 2	200 ± 44	59 ± 15
5007	[O III]	-0.040	579 ± 87	188 ± 49	586 ± 48	118 ± 7	590 ± 132	193 ± 56
5041	Si III]	-0.049	...	1:
5056	Si III]	-0.053	...	3:	9 ± 1	3:	33 ± 7	4 ± 1
5192	[Ar III]	-0.089	...	3 ± 1	...	12 ± 1	...	4 ± 1
5198	[N I]	-0.091	13 ± 1	8 ± 2	4:	...	5 ± 1	4 ± 1
5518	[Cl III]	-0.165	...	19 ± 5	24 ± 2	...	3:	13 ± 4
5538	[Cl III]	-0.170	...	7 ± 1	20 ± 2	10 ± 3
5755	[N II]	-0.213	3:	21 ± 6	10 ± 1	...	5 ± 1	21 ± 6
5876	He I	-0.235	14 ± 2	14 ± 3	13 ± 2	16 ± 1	9 ± 2	13 ± 4
6312	[S III]	-0.309	4:	5 ± 1	12 ± 1	3:
6563	H α	-0.348	286 ± 28	286 ± 75	287 ± 24	287 ± 24	287 ± 64	286 ± 83
6678	He I	-0.365	14 ± 2	15 ± 3	3 ± 1	28 ± 8
6716	[S II]	-0.371	10 ± 2	52 ± 12	17 ± 1	22 ± 2
6731	[S II]	-0.373	12 ± 1	40 ± 12	12 ± 4	73 ± 5	10 ± 2	17 ± 5
7065	He I	-0.419	2:	23 ± 5	2 ± 1
7136	[Ar III]	-0.428	8 ± 1	19 ± 4	17 ± 2	3:
7281	He I	-0.446	4:	...	5 ± 1	...
7320	[O II]	-0.451	3:	23 ± 5	13 ± 1	...	6 ± 2	...
7330	[O II]	-0.452	3:	30 ± 7	20 ± 2	5:	...	28 ± 8
7751	[Ar III]	-0.500	5 ± 1	...
8481	[Cl III]	-0.571	14 ± 1	4:
9069	[S III]	-0.617	20 ± 4	67 ± 17	31 ± 3	46 ± 3

Table 4
(Continued)

λ (Å)	Id.	f_λ	$I(\lambda)/I(\text{H}\beta), I(\text{H}\beta) = 100$					
			IRAS 08208+2816	IRAS 08339+6517	SBS 0926+606A	Arp 252	SBS 0948+532	Tol 9
9531	[S III]	-0.649	37 ± 3	127 ± 33	40 ± 3	15 ± 2	...	48 ± 14
$F(\text{H}\beta)^a$			13.7 ± 0.82	7.95 ± 1.1	4.5 ± 0.29	9.45 ± 0.26	4.96 ± 0.77	2.47 ± 0.37
$C(\text{H}\beta)$			0.13 ± 0.05	0.22 ± 0.05	0.10 ± 0.04	0.35 ± 0.05	0.57 ± 0.05	0.58 ± 0.06
EW_{abs}			3.4 ± 0.3	2.0 ± 0.5	0.07 ± 0.05	2.0 ± 0.5	0.5 ± 0.3	7.5 ± 0.5
$EW(\text{H}\alpha)^b$			554 ± 32	41 ± 6.6	917 ± 42	415 ± 26	634 ± 101	181 ± 31
$EW(\text{H}\beta)^b$			94 ± 5	8 ± 1.1	108 ± 5	135 ± 3	66 ± 10	22 ± 3
$EW([\text{O III}] 5007)^b$			610 ± 26	19 ± 3	720 ± 27	214 ± 8	419 ± 66	61 ± 10
$EW(\text{C III}] 1909)^b$			1:	1:	3:	7 ± 1	2:	1:

Notes.^a Units of $F(\text{H}\beta)$ are given in $\text{erg s}^{-1} \text{cm}^{-2} \times 10^{-15}$.^b Units of equivalent widths (EWs) are given in Å.

(ii) determine the fraction of O^{+2} , $X(\text{O}^{+2})$, (iii) use Figure 2 in Garnett et al. (1995) to obtain $X(\text{C}^{+2})$ and the ionization correction factor (ICF) for the unseen ions of carbon, $\text{ICF}(\text{C})$, and (iv) multiply the ionic abundance ratio $\text{C}^{+2}/\text{O}^{+2}$ by $\text{ICF}(\text{C})$ to obtain C/O :

$$\frac{\text{C}^{+2}}{\text{O}^{+2}} = 0.089 \times e^{\left(\frac{-1.09}{T_e/10^4}\right)} \times \frac{I(\text{C III}] 1909)}{I(\text{O III}] 1666)}, \quad (2)$$

$$X(\text{O}^{+2}) = \frac{\text{O}^{+2}}{\text{O}_{\text{total}}}, \quad (3)$$

$$\text{ICF}(\text{C}) = \left[\frac{X(\text{C}^{+2})}{X(\text{O}^{+2})} \right]^{-1}, \quad (4)$$

$$\text{C}/\text{O} = \frac{\text{C}^{+2}}{\text{O}^{+2}} \times \text{ICF}(\text{C}). \quad (5)$$

The constant 0.089 is the result of the product of the effective collision strengths between the two levels at electron temperatures below 20,000 K, $\Omega(1661, 1666)$ (Baluja et al. 1981) and $\Omega(1906, 1909)$ (Dufon et al. 1978), the statistical weight of the corresponding lower level, the excitation potential of the transition, and the number density of the ion under consideration. For the objects where we did not obtain a measurement for $\text{O III}] 1666 \text{ \AA}$, we used the intensity of $\text{O III}] 1661 \text{ \AA}$, which is possible because the transitions of both lines arise from the same level as explained in Garnett et al. (1995).

4.4.2. Total Abundances

The total gaseous abundances for O, Ne, and S were determined with the following equations and the ICFs given in Table 8:

$$\frac{N(\text{O})}{N(\text{H})} = \frac{N(\text{O}^+) + N(\text{O}^{++})}{N(\text{H}^+)}, \quad (6)$$

$$\begin{aligned} \frac{N(\text{Ne})}{N(\text{H})} &= \frac{N(\text{O}^+) + N(\text{O}^{++})}{N(\text{O}^{++})} \times \frac{N(\text{Ne}^{++})}{N(\text{H}^+)} \\ &= \text{ICF}(\text{Ne}) \times \frac{N(\text{Ne}^{++})}{N(\text{H}^+)}, \end{aligned} \quad (7)$$

and

$$\frac{N(\text{S})}{N(\text{H})} = \text{ICF}(\text{S}) \frac{N(\text{S}^+) + N(\text{S}^{++})}{N(\text{H}^+)}. \quad (8)$$

The resulting C abundances using the Garnett method, as well as the abundances of O, Ne, and S determined using the traditional analysis, are presented in Table 9. We refer to the standard assumption of a two-zone approximation to define the temperature structure of the object (with the electron temperature of $[\text{O III}]$ representing the high-ionization potential ions and that of $[\text{O II}]$ representing the low-ionization potential ions) and the use of such temperatures to determine ionic and total abundances as the traditional analysis. Since the low resolution of the spectra does not permit the $[\text{N II}]$ lines to be deblended from $\text{H}\alpha$, and the observed regions of the W-R galaxies of our sample match the center regions studied in the works of López-Sánchez et al., we adopted the N abundances of LSE08 and LSE10b and included these values in Table 9 for completeness. We used standard error propagation equations to determine the final uncertainties from our values and those given in the López-Sánchez et al. papers; nonetheless, additional sources of error may have been introduced.

4.4.3. Corrections to the Direct Method

The direct method has two essential shortcomings: (i) it depends on the capability to observe the weak auroral lines such as $[\text{O III}] 4363 \text{ \AA}$ and $[\text{S III}] 6312 \text{ \AA}$, which can prove quite difficult in distant objects, with high redshift, or objects that are intrinsically faint; and (ii) the temperature structure of the object is not taken into account, i.e., abundances can be significantly underestimated with the direct method due to large- and small-scale temperature inhomogeneities (Peimbert 1967; Peimbert & Costero 1969). There are several works in the literature that address this problem; for a review, see Peimbert & Peimbert (2011) and López-Sánchez et al. (2012).

In this work, we use the corrections proposed by Peña-Guerrero et al. (2012a). Peña-Guerrero et al. used a sample of 28 H II regions from the literature with the measured temperature inhomogeneity parameter, t^2 (Peimbert 1967), to derive a first approximation to the correction function of the O abundance determined with the auroral line $[\text{O III}] 4363 \text{ \AA}$, due to the thermal structure of the object as well as the fraction of

Table 5
Line Intensity Ratios for SBS 1054+365, POX 4, SBS 1319+579, SBS 1415+437, Tol 1457–262, and III Zw 107

λ (Å)	Id.	f_{λ}	$I(\lambda)/I(H\beta)$, $I(H\beta) = 100$					
			SBS 1054+365	POX 4	SBS 1319+579	SBS 1415+437	Tol 1457–262	III Zw 107
1661	O III]	1.124	2 ± 1	...	4 ± 1	2:
1666	O III]	1.122	5 ± 1	5:	10 ± 4	7 ± 3	7 ± 4	2 ± 1
1883	Si III]	1.187	20 ± 2	11 ± 1	4 ± 1	15 ± 2	5:	20 ± 1
1892	Si III]	1.197	18 ± 2	5 ± 1	5 ± 1	...	6:	...
1902	Fe II	1.208	20 ± 2	7 ± 1	31 ± 2	6 ± 2	1:	6:
1909	C III]	1.216	28 ± 8	23 ± 7	59 ± 18	4 ± 1	6 ± 1	13 ± 4
2322	[O III]	1.337	3 ± 1	7 ± 1	5 ± 1	4 ± 1	2:	8:
2471	[O II]	0.999	...	4 ± 1	2:	2 ± 1	1:	3:
2569	[Na VI]	0.858	9 ± 1	2:	4:
2661	[Al II]	0.764	1:	2:	...	3 ± 1	3:	4:
2670	[Al II]	0.757	4 ± 1	1:	2:	...	3:	1:
2783	[Mg V]	0.670	...	3:	...	2:	3:	2:
2854	[Ar IV]	0.623	4:	3:	4:
2868	[Ar IV]	0.615	4 ± 1	4:	...	5 ± 1	2:	2:
2871	[Na VI]	0.613	5 ± 1	3:	...	4:	...	5:
2928	[Mg V]	0.580	...	2:	1:	2:	1:	3:
2970	[Na VI]	0.558	2:	3:	1:	2:	4:	2:
3242	[Na IV]	0.440	...	10 ± 2	42 ± 3	159 ± 11
3362	[Na IV]	0.398	42 ± 3	3 ± 1	...	1:
3727	[O II]	0.291	64 ± 6	46 ± 4	79 ± 6	85 ± 5	130 ± 7	146 ± 10
3798	H 10+[S III]	0.272	9 ± 1	8 ± 1	4:	...
3869	[Ne III]	0.254	87 ± 6	38 ± 3	67 ± 5	14 ± 1	35 ± 3	19 ± 1
3889	H 8+He I	0.248	80 ± 6	17 ± 1	16 ± 1	36 ± 2
3967	[Ne III]	0.229	56 ± 4	23 ± 5	30 ± 2	32 ± 2	24 ± 2	17 ± 1
4069	[S II]	0.203	1:	11:	...	7:	2:	20 ± 1
4076	[S II]	0.202	6:	...	2:	8:	3:	12 ± 1
4102	H δ	0.195	28 ± 4	19 ± 2	27 ± 2	...	25 ± 2	25 ± 2
4125	[Fe II]	0.190	...	6 ± 1	3:	...	6 ± 1	...
4340	H γ	0.137	47 ± 3	47 ± 3	49 ± 6	48 ± 3	48 ± 3	47 ± 3
4363	[O III]	0.132	9 ± 1	9 ± 1	11:	6:	8 ± 1	4:
4563	Mg I]	0.081	...	2:	...	8 ± 1	...	6 ± 1
4571	Mg I]	0.079	11 ± 8	...	1:	...	6 ± 1	6 ± 1
4658	[Fe III]	0.056	...	5:	2:	...	4:	...
4686	He II	0.048	<2	<3	<2	...	1:	...
4702	[Fe III]	0.044	...	5 ± 1	4:	9:
4711	[Ar IV]+He I	0.041	...	8 ± 1	4:	14 ± 1	...	19 ± 2
4740	[Ar IV]	0.034	...	2 ± 1	4:	...
4755	[Fe III]	0.029	41 ± 3	4:
4861	H β	0.000	100 ± 7	100 ± 7	100 ± 7	100 ± 7	100 ± 6	100 ± 7
4922	He I	-0.017	...	7:	1:	3:	11 ± 1	2:
4959	[O III]	-0.027	217 ± 21	199 ± 17	301 ± 58	100 ± 8	141 ± 14	159 ± 12
5007	[O III]	-0.040	647 ± 48	585 ± 42	911 ± 182	272 ± 19	411 ± 29	460 ± 31
5041	Si III]	-0.049	32 ± 2	6:	2:	12 ± 1	1:	...
5192	[Ar III]	-0.089	32 ± 2	3:	1:	6:	2:	...
5198	[N I]	-0.091	32 ± 2	4:	1:	15 ± 5	2:	18 ± 7
5518	[Cl III]	-0.165	...	2:	2:	...	8 ± 1	...
5538	[Cl III]	-0.170	...	7:	...	27 ± 2	16 ± 1	30 ± 2
5755	[N II]	-0.213	6:	6:	13 ± 1	4:
5876	He I	-0.235	12 ± 2	12 ± 1	14 ± 1	9 ± 1	15 ± 2	30 ± 2
6312	[S III]	-0.309	39 ± 3	...	1:	28 ± 2	5:	20 ± 2
6371	[S III]	-0.318	11 ± 1	6:	4:
6563	H α	-0.348	280 ± 20	287 ± 18	286 ± 21	274 ± 14	287 ± 21	286 ± 29
6678	He I	-0.365	...	7 ± 1	...	11 ± 1
6716	[S II]	-0.371	4:
7065	He I	-0.419	2:	...	4:	2:	4:	7:
7136	[Ar III]	-0.428	...	5:	4:	...	6:	6:
7281	He I	-0.446	44 ± 3	34 ± 2
7320	[O II]	-0.451	6:	7:	2:	...	12 ± 1	6:
7751	[Ar III]	-0.500	...	2:	2:	19 ± 1	4:	20 ± 1
8481	[Cl III]	-0.571	3:	...	1:	...
9069	[S III]	-0.617	84 ± 6	5:	8 ± 1	...	7:	22 ± 2
9531	[S III]	-0.649	...	13 ± 1	13 ± 1

Table 5
(Continued)

λ (Å)	Id.	f_{λ}	$I(\lambda)/I(H\beta)$, $I(H\beta) = 100$					
			SBS 1054+365	POX 4	SBS 1319+579	SBS 1415+437	Tol 1457–262	III Zw 107
$F(H\beta)^a$			2.23 ± 0.11	12.32 ± 0.62	10.6 ± 0.54	3.20 ± 0.16	8.55 ± 0.39	6.1 ± 0.29
$C(H\beta)$			0.01 ± 0.05	0.22 ± 0.05	0.01 ± 0.05	0.00 ± 0.05	0.04 ± 0.06	0.23 ± 0.05
EW_{abs}			0.8 ± 0.3	2.0 ± 0.4	2.0 ± 0.5	1.0 ± 0.5	0.04 ± 0.06	0.23 ± 0.05
$EW(H\alpha)^b$			1096 ± 48	611 ± 28	1690 ± 68	337 ± 18	844 ± 59	714 ± 61
$EW(H\beta)^b$			98 ± 4	101 ± 5	315 ± 5	96 ± 4	126 ± 5	310 ± 14
$EW([O III] 5007)^b$			659 ± 26	719 ± 50	3235 ± 275	294 ± 27	517 ± 22	2397 ± 107
$EW(C III] 1909)^b$			2:	1:	8 ± 3	1:	2:	3 ± 1

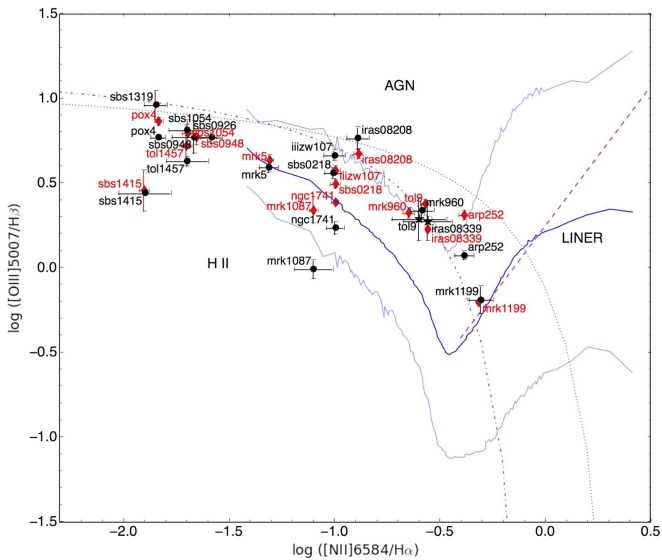
Notes.^a Units of $F(H\beta)$ are given in $\text{erg s}^{-1} \text{cm}^{-2} \times 10^{-15}$.^b Units of equivalent widths (EWs) are given in Å.

Figure 5. Behavior of $\log([O III]/H\beta)$ to $\log([N II]/H\alpha)$ from observed W-R galaxy sample. Filled circles represent the objects for which we obtained a T_e measurement, the black stars represent those objects for which we used T_e from the literature, and the blue triangles represent reference points of objects studied in great detail: data for I Zw 18 was taken from Leboutteiller et al. (2013), data for the Sun was taken from Asplund et al. (2009), data for 30 Dor was taken from Peimbert (2003), and data for Orion was taken from Esteban et al. (2004). The red diamonds represent the data from López-Sánchez et al. (2004a) for NGC 1741, López-Sánchez et al. (2004b) for Mrk 1087, López-Sánchez et al. (2008) for IRAS 08339+6517, and from LSE09 for the rest of the objects. The blue solid curve is the SDSS outline for the median of $\log([O III]/H\beta)$ to $\log([N II]/H\alpha)$, while the light blue curves are the same $\pm 3\sigma$, respectively. The green dotted line represents the theoretical upper star formation limit from Kewley et al. (2001), the dashed-dotted magenta line represents the lower limit for AGNs from Kauffmann et al. (2003), and the dashed magenta line represents the division between AGNs and LINERs from Kauffmann et al. (2003).

oxygen depleted into dust grains. The authors refer to this relation as the Corrected Auroral Line Method (CALM): $12 + \log(O/H)_{\text{CALM}} = 1.0825 \times (O/H)_{\text{Direct Method}} - 0.375$. The authors then applied this correction function to the relations given by Pilyugin & Thuan (2005) for the upper and lower branches of the $12 + \log(O/H)$ versus R_{23} diagram, where the strong line metallicity indicator $R_{23} = I([O II] 3727) + I([O III] 4959 + 5007) / I(H\beta)$, obtaining a strong line method that accounts for the thermal structure, dust depletion, and the ionization structure of the object. Peña-Guerrero et al. refer to this technique as the Recalibrated R_{23} Method (RRM).

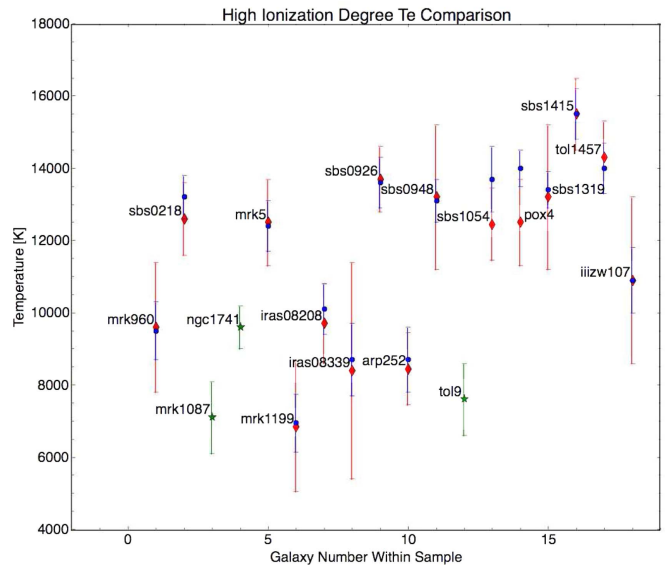


Figure 6. Comparison of our high-ionization zone temperatures vs. those presented in LSE08, LSE09, LSE10a, LSE10b. Red diamonds represent our measurements, blue circles represent the data from LSE, and the green stars represent the temperatures from López-Sánchez & Esteban that we adopted.

When the R_{23} abundance falls in between 8.29 and 8.55 (often called the degeneracy zone), the set of equations and values to use is undefined. In such cases, to know which set of equations is appropriate, one can either use another metallicity indicator, such as $[N II]/H\alpha$ (Storchi-Bergmann et al. 1994), that increases linearly with $12 + \log(O/H)$, or use an indicator of the hardness of the ionizing radiation. Pilyugin (2000) proposed the excitation index, $P = I([O III] 4959 + 5007) / (I([O II] 3727) + I([O III] 4959 + 5007))$, to be one such indicator. Peña-Guerrero et al. (2012a) introduced the Oxygen Ionization Degree, $OID = O^{++} / (O^{+} + O^{++})$, which is a quantity equivalent to P .

We used the the CALM and RRM methods to calculate the oxygen abundances corrected for dust depletion presented in Table 10. Column 1 is the galaxy name, and columns 2 and 3 are, respectively, the carbon abundances determined through the method described in Garnett et al. (1995) and corrected for dust depletion. Columns 4 through 7 are, respectively, the oxygen abundances determined through the direct method, corrected for dust depletion as described in Peimbert & Peimbert (2010), CALM as described in Peña-Guerrero et al. (2012a), and RRM as described in Peña-Guerrero et al. (2012a). Column

Table 6
Physical Properties of the Observed Regions in Our W-R Galaxy Sample

Galaxy Name	T_e High ionization zone (K)	T_e Low ionization zone (K)	Electron Density (cm^{-3})
Mrk 960	9600 ± 1800	9700 ± 1800^b	<100
SBS 0218+003	$12,600 \pm 1000$	$11,800 \pm 1000^b$	<600
Mrk 1087	7100 ± 1000^a	8000 ± 1100^a	120 ± 90
NGC 1741	9600 ± 600^a	9600 ± 600^a	<100
Mrk 5	$12,500 \pm 1200$	$11,750 \pm 1200^b$	<200
Mrk 1199	6850 ± 1800	7800 ± 1800^b	<400
IRAS 08208+2816	9700 ± 1100	9800 ± 1000^b	<300
IRAS 08339+6517	8400 ± 3000	8900 ± 3050^b	<100
SBS 0926+606A	$13,700 \pm 900$	$12,600 \pm 900^b$	<100
Arp 252	8450 ± 1000	8900 ± 1100^b	<400
SBS 0948+532	$13,200 \pm 2000$	$12,200 \pm 2000^b$	250 ± 80^a
Tol 9	7600 ± 1000^a	8300 ± 700^a	180 ± 60^a
SBS 1054+365	$12,450 \pm 1000$	$11,700 \pm 1000^b$	100^a
POX 4	$12,500 \pm 1200$	$11,750 \pm 1200^b$	250 ± 80^a
SBS 1319+579	$13,200 \pm 2000$	$12,200 \pm 2000^b$	<200
SBS 1415+437	$15,600 \pm 1000$	$13,850 \pm 800^b$	<100 ^a
Tol 1457–262	$14,300 \pm 1000$	$13,000 \pm 900^b$	200 ± 80^a
III Zw 107	$10,900 \pm 2300$	$10,500 \pm 2300^b$	200 ± 60^a

Notes.

^a Adopted from values determined by LSE09 and LSE10b.

^b Values obtained using Equation (2) in Garnett (1992).

8 is the strong line metallicity indicator R_{23} as defined by Pagel et al. (1979). Column 9 shows the oxygen excitation ratio, P , as defined by Pilyugin (2000), and column 10 presents the OID, as defined by Peña-Guerrero et al. (2012a).

5. Discussion

5.1. C/O versus O/H Diagram

The resulting C/O ratios and the gaseous oxygen abundance measured from our STIS observations are plotted in Figure 7. It is interesting to note that the C/O in this figure and Figure 14 in Henry & Nava (2006), as well as the N/O in Figure 12 of Nava et al. (2006), shows an increase with respect to O/H starting around $12+\log(\text{O}/\text{H}) \sim 8.2$. This behavior is likely due to the contribution to C and N by intermediate-mass stars, which in turn implies that both carbon and nitrogen could be mainly produced in the same stars; however, it is important to note that the slopes of the previously mentioned figures are quite different.

The trend in Figure 7 resembles that of the equivalent figure in Garnett et al. (1995), which shows an apparent increase in C/O with increasing O/H. Garnett et al. found that a good fit to their data was a power law of the form $\log(\text{C}/\text{O}) = A + B \log(\text{O}/\text{H})$, with $A = 1.01 \pm 0.39$ and $B = 0.43 \pm 0.09$ for the abundance range $7.3 \leq 12+\log(\text{O}/\text{H}) \leq 8.7$. In this work, we find that a power law may not be the best fit for our data. Though there is also an apparent increase of C/O with respect to O/H in our data, the behavior does not follow a specific curve, particularly when taking into account I Zw 18 (Lebouteiller et al. 2013), included as reference point along with 30 Doradus (Peimbert 2003), Orion (Esteban et al. 2004), and the Sun (Asplund et al. 2009). Nonetheless, we consider that there is a section of the diagram ($12+\log(\text{O}/\text{H}) > 7.5$ or $[\text{O}/\text{H}] > 1.23$) whose behavior could be described by the

linear function

$$\log(\text{C}/\text{O}) = m \log(\text{O}/\text{H}) + b, \quad (9)$$

where $m = 1.8 \pm 0.4$, $b = -14.9 \pm 2.9$, and the correlation coefficient is 0.78. The quantity $\log(\text{O}/\text{H})$ is given in units of $12+\log(\text{O}/\text{H})$. In this figure, we also present the linear fit of Garnett et al. (1995) and the data from Berg et al. (2016), as well as their literature points, for comparison. The data from the low-metallicity high-ionization H II regions in the Berg et al. work, as well as their literature data, seem to agree better with the fit of Garnett. This could suggest that there is a dependence of the slope of C/O versus O/H with respect to the IMF, since our sample has objects with a top-heavy IMF whereas the Berg et al. sample does not. This, in turn, could imply that massive stars contribute more efficiently to the production of C in objects with a top-heavy IMF and with metallicity $12+\log(\text{O}/\text{H}) \gtrsim 8.0$.

Assuming a simple chemical evolution model with instantaneous recycling, the expected outcome for this plot would be a constant value for C/O. Such constant behavior would imply that either both C and O are primary elements, or that O is primary and C secondary but with $\text{C}/\text{O} \propto \text{O}/\text{H}$. If we consider that only the primary carbon ‘‘pollutes’’ the ISM, an increase of C/O with increasing O/H would imply one or both of the following: (i) the instantaneous recycling approximation does not hold for both C and O, and (ii) the yield of C varies with metallicity.

5.2. Behavior of C/N

There does not seem to be a simple correlation between $\log(\text{C}/\text{N})$ and $12+\log(\text{O}/\text{H})$ in the data if we consider the presence of I Zw 18, (Figure 8). Nonetheless, omitting I Zw 18, there is a part of the diagram that could be described by a linear fit. We perform a linear fit to our data to describe such part of

Table 7
Ionic Chemical Composition of the Observed Regions in Our W-R Galaxy Sample^a

Galaxy Name	He ⁺ /H ⁺	O ⁺⁺ /H ⁺	O ⁺ /H ⁺	C ⁺⁺ /H ⁺	Ne ⁺⁺	/H ⁺ S ⁺⁺ /H ⁺	S ⁺ /H ⁺
Mrk 960	10.99:	7.96 ± 0.17	7.59 ± 0.22	6.81 ± 0.20	7.43 ± 0.20	6.60 ± 0.20	6.53 ± 0.20
SBS 0218+003	10.92:	7.80 ± 0.06	7.53 ± 0.08	6.84 ± 0.19	7.08 ± 0.10	5.95 ± 0.05	6.07 ± 0.02
Mrk 1087	10.88:	8.07 ± 0.14	8.13 ± 0.19	8.58 ± 0.19	7.92 ± 0.19	...	6.67 ± 0.22
NGC 1741	10.97:	7.86 ± 0.09	8.02 ± 0.12	7.83 ± 0.09	7.29 ± 0.15	6.39 ± 0.10	6.04 ± 0.13
Mrk 5	10.87:	7.84 ± 0.09	7.33 ± 0.12	6.95 ± 0.19	7.42 ± 0.14	6.22 ± 0.09	5.81 ± 0.12
Mrk 1199	10.83:	8.04 ± 0.19	8.02 ± 0.25	8.89 ± 0.23	7.44 ± 0.25	6.52 ± 0.22	6.32 ± 0.29
IRAS 08208+2816	11.01:	8.37 ± 0.10	7.53 ± 0.13	7.54 ± 0.15	7.76 ± 0.12	6.39 ± 0.11	5.85 ± 0.15
IRAS 08339+6517	10.98:	8.11 ± 0.06	8.05 ± 0.09	8.26 ± 0.19	8.02 ± 0.09	7.06 ± 0.08	5.76 ± 0.06
SBS 0926+606A	10.99:	7.91 ± 0.08	7.20 ± 0.15	6.48 ± 0.10	6.74 ± 0.13	6.14 ± 0.13	5.48 ± 0.17
Arp 252	11.05:	7.90 ± 0.13	7.74 ± 0.17	7.97 ± 0.19	7.67 ± 0.13	6.11 ± 0.12	6.75 ± 0.16
SBS 0948+532	10.90:	7.96 ± 0.10	7.21 ± 0.10	7.33 ± 0.15	7.37 ± 0.10	...	5.57:
Tol 9	10.96:	8.30 ± 0.13	8.25 ± 0.17	8.56 ± 0.21	7.82 ± 0.16	6.74 ± 0.10	6.21 ± 0.08
SBS 1054+365	10.96:	8.07 ± 0.08	7.07 ± 0.11	7.38 ± 0.11	7.62 ± 0.12	...	5.96:
POX 4	10.94:	8.02 ± 0.05	6.97 ± 0.07	7.29 ± 0.08	7.25 ± 0.07	5.72 ± 0.10	5.18 ± 0.14
SBS 1319+579	10.95:	8.14 ± 0.08	7.17 ± 0.11	6.52 ± 0.09	7.42 ± 0.09	5.67 ± 0.14	5.21 ± 0.19
SBS 1415+437	10.83:	7.44 ± 0.06	6.94 ± 0.09	5.93 ± 0.17	6.54 ± 0.08	...	5.91:
Tol 1457–262	11.01:	7.71 ± 0.14	7.24 ± 0.19	6.41 ± 0.30	7.04 ± 0.12	...	5.41:
III Zw 107	11.05:	8.10 ± 0.09	7.62 ± 0.12	7.47 ± 0.18	7.18 ± 0.12	...	6.14:

Note.

^a In units of $12 + \log n(X^{i+})/n(H^+)$. Gaseous content only.

Table 8
Ionization Correction Factors

Galaxy Name	C ^a	Ne ^b	S ^c
Mrk 960	1.25 ± 0.45	1.42 ± 0.15	1.12 ± 0.20
SBS 0218+003	1.02 ± 0.45	1.53 ± 0.13	1.12 ± 0.20
Mrk 1087	1.05 ± 0.45	2.15 ± 0.23	1.12 ± 0.20
NGC 1741	1.00 ± 0.45	2.45 ± 0.21	1.12 ± 0.20
Mrk 5	1.05 ± 0.20	1.30 ± 0.10	1.26 ± 0.20
Mrk 1199	1.05 ± 0.45	1.97 ± 0.31	1.12 ± 0.20
IRAS 08208+2816	1.00 ± 0.20	1.14 ± 0.11	1.78 ± 0.20
IRAS 08339+6517	1.05 ± 0.45	1.86 ± 0.40	1.12 ± 0.20
SBS 0926+606A	1.25 ± 0.20	1.16 ± 0.13	1.78 ± 0.20
Arp 252	1.02 ± 0.45	1.68 ± 0.08	1.12 ± 0.20
SBS 0948+532	1.11 ± 0.20	1.18 ± 0.39	1.26 ± 0.20
Tol 9	0.95 ± 0.45	1.88 ± 0.45	1.12 ± 0.20
SBS 1054+365	1.25 ± 0.20	1.10 ± 0.13	1.78 ± 0.20
POX 4	1.18 ± 0.20	1.09 ± 0.08	1.78 ± 0.20
SBS 1319+579	1.25 ± 0.20	1.11 ± 0.08	1.78 ± 0.20
SBS 1415+437	1.00 ± 0.20	1.31 ± 0.06	1.26 ± 0.20
Tol 1457–262	1.11 ± 0.45	1.34 ± 0.10	1.26 ± 0.20
III Zw 107	1.09 ± 0.45	1.33 ± 0.10	1.26 ± 0.20

Notes.

^a Obtained from Garnett et al. (1995).

^b Obtained from Peimbert & Costero (1969).

^c Obtained from Garnett (1989).

the diagram and obtain

$$\log(C/N) = m \log(O/H) + b, \quad (10)$$

where $m = 0.8 \pm 0.3$, $b = -6.0 \pm 2.6$, and the correlation coefficient is 0.60. The quantity $\log(O/H)$ is given in units of $12 + \log(O/H)$. If true, this linear increase of C/N with increasing O/H would imply that there is an additional contribution of C at higher metallicities, and an additional contribution of N at lower metallicities. This figure contrasts with Figure 6(b) of Berg et al. (2016), in which the authors find a relatively constant behavior. The difference in our findings versus those in Berg et al. could

be due to the physical differences of the samples used. They used low-metallicity and high-ionization H II regions in dwarf galaxies while we have a range of both low- and high-ionization degrees and metallicities. Furthermore, our sample is composed of top-heavy IMF objects. Hence, the difference in the figures could suggest that the production of carbon and nitrogen in objects with such an IMF has a strong contribution from massive stars, and that these stars favor the production of C over N at metallicities higher than $12 + \log(O/H) \gtrsim 8.0$. Figure 8 shows significant scatter, and the linear behavior seems to be only true for objects in our sample, again indicating that the origin of C and N production is not homogeneous for objects with different IMFs.

The C/N to O/H figure in Garnett et al. (1995) appears to resemble a doubled-valued curve similar to a negative parabola; nonetheless, they describe it as not showing a clear correlation. In our observations, we did not have the resolution to separate the [N II] emission lines from H α ; therefore, we adopted the nitrogen abundances derived by LSE08 and LSE10b. We used standard error propagation equations to combine our uncertainties with those given in the López-Sánchez & Esteban papers. The reader should, however, be aware that additional sources of error may have been introduced. Figure 9 shows a clear correlation for our STIS data; however, it is again evident that such behavior is not the same for objects with a different IMF. We find that our data are consistent with a linear fit of the form

$$\log(C/N) = m \log(C/H) + b, \quad (11)$$

with $b = -2.1 \pm 0.7$, slope $m = 0.4 \pm 0.1$, and a correlation coefficient of 0.68.

To determine if there is any relation between $\log(C/N)$ and $\log(N/O)$, we plotted these quantities in Figure 10. Though the correlation coefficient is very close to zero, we ran an MCMC model of a linear fit to our STIS data in order to get a sense of the possible fits. We used 100 walkers and did 500 runs. As expected, we found that about half of the fitted lines show a

Table 9
Total Chemical Composition of the Observed Regions in Our W-R Galaxy Sample^a

Galaxy Name	O	C	N ^b	Ne	S	C/O	N/O	Ne/O	S/O
Mrk 960	8.11 ± 0.19	7.49 ± 0.20	7.34 ± 0.10	7.59 ± 0.20	6.90 ± 0.2	-0.62 ± 0.28	-1.39 ± 0.28	-0.52 ± 0.28	-1.21 ± 0.28
SBS 0218+003	7.99 ± 0.06	6.82 ± 0.19	6.84 ± 0.06	7.27 ± 0.10	6.36 ± 0.09	-1.44 ± 0.20	-1.15 ± 0.08	-0.72 ± 0.12	-1.63 ± 0.11
Mrk 1087	8.41 ± 0.16	8.60 ± 0.19	7.76 ± 0.16	8.25 ± 0.19	6.79 ± 0.19 ^b	0.19 ± 0.25	-0.65 ± 0.23	-0.16 ± 0.25	-1.62 ± 0.25
NGC 1741	8.25 ± 0.10	7.83 ± 0.20	7.10 ± 0.09	7.68 ± 0.15	6.60 ± 0.11	-0.42 ± 0.22	-1.15 ± 0.13	-0.57 ± 0.18	-1.65 ± 0.15
Mrk 5	7.98 ± 0.10	6.97 ± 0.19	6.72 ± 0.09	7.54 ± 0.14	6.46 ± 0.10	-0.99 ± 0.21	-1.24 ± 0.13	-0.42 ± 0.17	-1.50 ± 0.14
Mrk 1199	8.33 ± 0.22	8.91 ± 0.23	7.90 ± 0.18	7.73 ± 0.25	6.78 ± 0.25	0.58 ± 0.32	-0.43 ± 0.28	-0.70 ± 0.33	-1.55 ± 0.33
IRAS 08208+2816	8.43 ± 0.11	7.54 ± 0.15	7.44 ± 0.10	7.82 ± 0.12	6.75 ± 0.13	-0.89 ± 0.19	-0.99 ± 0.15	-0.61 ± 0.16	-1.68 ± 0.17
IRAS 08339+6517	8.42 ± 0.07	8.27 ± 0.19	7.51 ± 0.10	8.25 ± 0.09	7.13 ± 0.12	-0.15 ± 0.20	-0.91 ± 0.07	-0.17 ± 0.11	-1.29 ± 0.14
SBS 0926+606A	7.98 ± 0.09	7.29 ± 0.15	6.48 ± 0.10	6.81 ± 0.13	6.49 ± 0.15	-0.69 ± 0.17	-1.50 ± 0.13	-1.17 ± 0.16	-1.49 ± 0.13
Arp 252	8.13 ± 0.15	7.98 ± 0.19	7.71 ± 0.08	7.89 ± 0.13	6.89 ± 0.14	-0.15 ± 0.24	-0.42 ± 0.17	-0.24 ± 0.20	-1.24 ± 0.17
SBS 0948+532	8.03 ± 0.09	7.33 ± 0.15	6.61 ± 0.07	7.44 ± 0.10	6.34 ± 0.11 ^b	-0.70 ± 0.17	-1.42 ± 0.11	-0.59 ± 0.13	-1.69 ± 0.14
Tol 9	8.58 ± 0.15	8.54 ± 0.21	7.80 ± 0.14	8.10 ± 0.16	6.96 ± 0.16	-0.04 ± 0.26	-0.78 ± 0.21	-0.48 ± 0.22	-1.62 ± 0.17
SBS 1054+365	8.11 ± 0.09	7.48 ± 0.11	6.59 ± 0.09	7.66 ± 0.12	6.21 ± 0.18 ^b	-0.63 ± 0.14	-1.52 ± 0.13	-0.45 ± 0.15	-1.90 ± 0.20
POX 4	8.06 ± 0.05	7.36 ± 0.09	6.50 ± 0.06	7.29 ± 0.07	6.08 ± 0.12	-0.69 ± 0.09	-1.55 ± 0.08	-0.76 ± 0.09	-1.97 ± 0.13
SBS 1319+579	8.19 ± 0.09	7.66 ± 0.17	6.52 ± 0.09	7.46 ± 0.09	6.05 ± 0.16	-0.53 ± 0.19	-1.67 ± 0.13	-0.73 ± 0.19	-2.14 ± 0.13
SBS 1415+437	7.56 ± 0.07	5.93 ± 0.17	6.04 ± 0.11	6.66 ± 0.08	5.89 ± 0.13 ^b	-1.64 ± 0.22	-1.52 ± 0.13	-0.90 ± 0.11	-1.67 ± 0.15
Tol 1457-262	7.83 ± 0.16	6.40 ± 0.30	6.48 ± 0.09	7.16 ± 0.12	6.18 ± 0.10 ^b	-1.43 ± 0.31	-1.35 ± 0.13	-0.67 ± 0.15	-1.65 ± 0.13
III Zw 107	8.22 ± 0.10	7.50 ± 0.18	7.07 ± 0.08	7.30 ± 0.12	6.42 ± 0.13 ^b	-0.72 ± 0.21	-1.15 ± 0.13	-0.92 ± 0.16	-1.80 ± 0.16

Notes.

^a In units of $12 + \log n(X)/n(H)$. Gaseous content only.

^b Values taken from the work of [LSE09](#) and [LSE10b](#). These abundances are presented for completeness and for comparison with modeled He abundances.

Table 10
Corrections to the Direct Method Carbon and Oxygen Abundances^a

Galaxy Name	12+log(C/H)		12+log(O/H)				R_{23}	P	OID ^h
	This Work ^b	+Dust ^c	Direct ^d	+Dust ^e	CALM ^f	RRM ^g			
Mrk 960	7.49 ± 0.20	7.69	8.11 ± 0.19	8.21	8.41 ± 0.21	<i>dz</i>	3.89 ± 0.19	0.75 ± 0.08	0.70 ± 0.10
SBS 0218+003	6.82 ± 0.19	7.02	7.99 ± 0.06	8.09	8.27 ± 0.06	8.10 ± 0.12	6.71 ± 0.32	0.72 ± 0.09	0.65 ± 0.10
Mrk 1087	8.81 ± 0.19	9.01	8.42 ± 0.16	8.52	8.72 ± 0.17	8.92 ± 0.19	2.79 ± 0.19	0.47 ± 0.09	0.47 ± 0.10
NGC 1741	7.83 ± 0.20	8.05	8.25 ± 0.10	8.35	8.55 ± 0.11	8.68 ± 0.10	4.97 ± 0.28	0.46 ± 0.10	0.41 ± 0.12
Mrk 5	6.97 ± 0.19	7.17	7.96 ± 0.10	8.06	8.24 ± 0.11	7.89 ± 0.12	6.32 ± 0.31	0.82 ± 0.10	0.72 ± 0.11
Mrk 1199	8.91 ± 0.23	9.11	8.33 ± 0.22	8.44	8.64 ± 0.24	9.05 ± 0.24	1.79 ± 0.22	0.48 ± 0.09	0.51 ± 0.10
IRAS 08208+2816	7.54 ± 0.15	7.74	8.43 ± 0.11	8.54	8.75 ± 0.12	8.66 ± 0.12	8.61 ± 0.51	0.90 ± 0.11	0.87 ± 0.12
IRAS 08339+6517	8.27 ± 0.19	8.47	8.42 ± 0.07	8.53	8.74 ± 0.08	8.90 ± 0.09	3.99 ± 0.40	0.63 ± 0.16	0.67 ± 0.19
SBS 0926+606A	7.29 ± 0.15	7.49	7.98 ± 0.09	8.08	8.21 ± 0.10	7.98 ± 0.15	7.66 ± 0.42	0.88 ± 0.10	0.84 ± 0.11
Arp 252	7.98 ± 0.19	8.18	8.13 ± 0.15	8.24	8.43 ± 0.16	<i>dz</i>	2.46 ± 0.10	0.64 ± 0.10	0.59 ± 0.12
SBS 0948+532	7.33 ± 0.15	7.53	8.03 ± 0.09	8.13	8.32 ± 0.09	<i>dz</i>	8.87 ± 0.95	0.89 ± 0.21	0.85 ± 0.23
Tol 9	8.54 ± 0.21	8.74	8.58 ± 0.15	8.69	8.98 ± 0.17	8.73 ± 0.18	5.45 ± 0.90	0.58 ± 0.17	0.56 ± 0.17
SBS 1054+365	7.48 ± 0.11	7.68	8.11 ± 0.09	8.21	8.35 ± 0.10	<i>dz</i>	9.25 ± 0.52	0.93 ± 0.13	0.91 ± 0.15
POX 4	7.36 ± 0.09	7.56	8.06 ± 0.05	8.15	8.34 ± 0.05	<i>dz</i>	8.27 ± 0.05	0.94 ± 0.10	0.92 ± 0.12
SBS 1319+579	7.66 ± 0.17	7.86	8.19 ± 0.09	8.29	8.49 ± 0.10	<i>dz</i>	12.94 ± 0.70	0.94 ± 0.10	0.90 ± 0.12
SBS 1415+437	5.93 ± 0.17	6.13	7.56 ± 0.07	7.66	7.81 ± 0.09	7.68 ± 0.11	4.85 ± 0.67	0.68 ± 0.15	0.76 ± 0.16
Tol 1457–262	6.40 ± 0.30	6.60	7.83 ± 0.16	7.93	8.11 ± 0.17	7.98 ± 0.17	6.78 ± 0.31	0.81 ± 0.10	0.75 ± 0.12
III Zw 107	7.50 ± 0.18	7.70	8.22 ± 0.10	8.32	8.53 ± 0.11	<i>dz</i>	7.75 ± 0.26	0.81 ± 0.10	0.75 ± 0.12

Notes.

^a In units of 12+log $n(X)/n(H)$.

^b Abundances determined using the method described in Garnett et al. (1995).

^c Consistent correction of 0.2 dex as discussed in the text, Section 5.4. Uncertainties remain the same.

^d Determined from the direct method, that is, using $T_e[OIII]$ as measured from the line ratio $[O III] 4363/5007$.

^e Using the metallicity-dependent correction from 0.08 to 0.12 dex according to Peimbert & Peimbert (2010). Uncertainties remain the same as with the direct method.

^f Corrected Auroral Line Method as described in Peña-Guerrero et al. (2012a).

^g Recalibrated R_{23} Method as described in Peña-Guerrero et al. (2012a). An abundance in the degeneracy zone is marked as *dz*.

^h Oxygen Ionization Degree as defined in Peña-Guerrero et al. (2012a): $OID = O^{++}/(O^+ + O^{++})$.

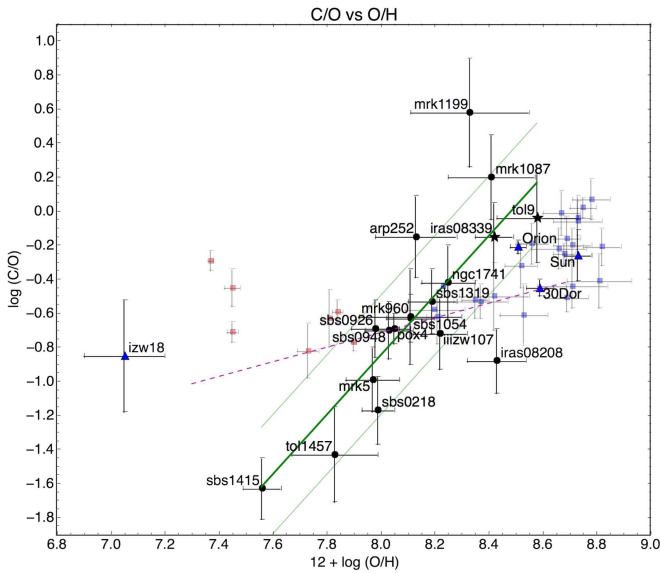


Figure 7. Behavior of C/O vs. 12+log(O/H). Filled circles represent the objects for which we obtained a T_e measurement; the black stars represent those objects for which we used T_e from the literature; the light red squares represent the data from Berg et al. (2016); the light blue squares represent additional data presented in Berg et al. (2016) taken from Esteban et al. (2002, 2004, 2014), García-Rojas & Esteban (2007), and López-Sánchez et al. (2007); and the blue triangles represent literature reference points: data for I Zw 18 were taken from Lehouetteiller et al. (2013), data for the Sun were taken from Asplund et al. (2009), data for 30 Dor were taken from Peimbert (2003), and data for Orion were taken from Esteban et al. (2004). The green line represents a linear fit to our data, and the lighter green lines represent the errors of the fit; the dashed magenta line represent the fit proposed in Garnett et al. (1995).

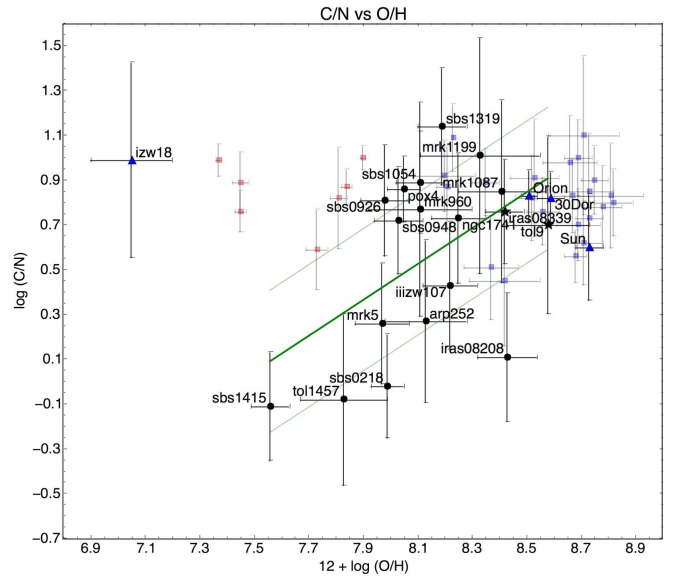


Figure 8. Behavior of log(C/N) vs. 12+log(O/H) as obtained from observations. Symbols are the same as in Figure 7. The green line represents a linear fit to our data, and the lighter green lines represent the errors of the fit.

positive slope and half a negative slope. As an experiment, we took the average of the fits with a negative slope to obtain a best estimate for parameters m and b , and similarly, we obtained another equation from the positive slopes. We then used these equations to determine the carbon abundances for the objects in the literature that have measurements of nitrogen and oxygen

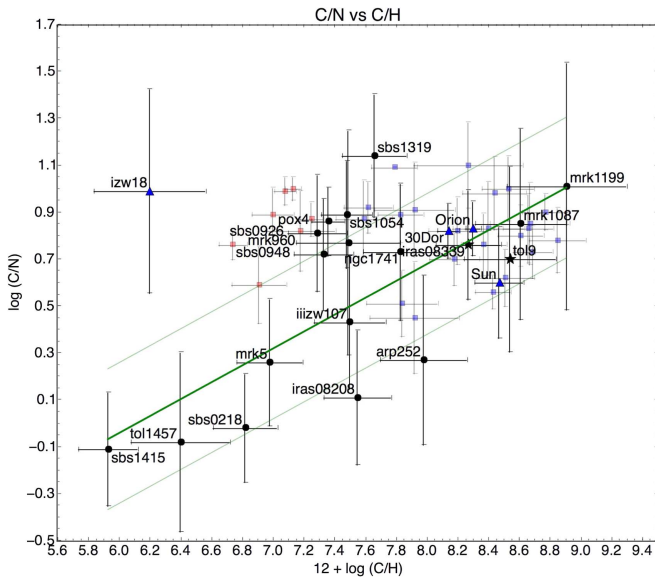


Figure 9. Behavior of C/N vs. $12 + \log(C/H)$ as obtained from observations. Symbols are the same as in Figure 7. The green line is the best fit found through an MCMC method, and the lighter green lines represent the errors of the fit. Only the objects in the sample were used for the fit; the Sun, Orion, 30 Dor, and I Zw 18 are also plotted as reference points.

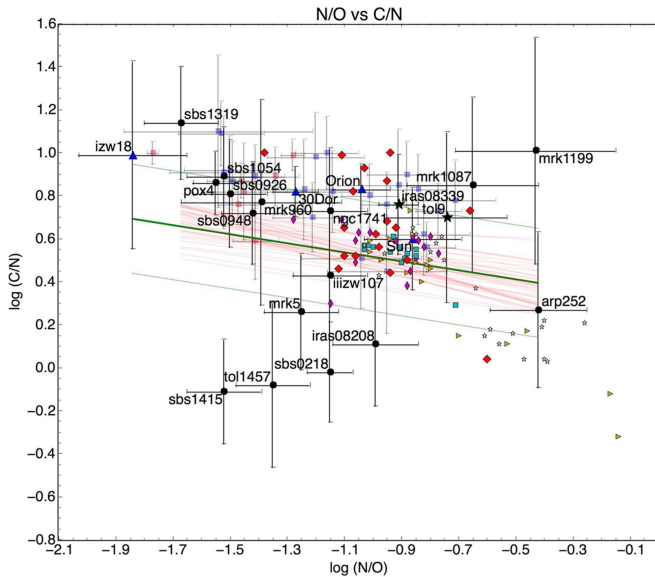


Figure 10. Behavior of $\log(C/N)$ vs. $\log(N/O)$ as obtained from observations. Filled circles, black stars, and blue triangles represent the same as in Figure 7; the light red squares represent the data from Berg et al. (2016); and the light blue squares represent additional data presented in Berg et al. (2016) taken from Esteban et al. (2002, 2004, 2014), García-Rojas & Esteban (2007), and López-Sánchez et al. (2007). The following are data for disk stars: magenta diamonds are data from Cunha & Lambert (1994), cyan squares from Nieva & Simon-Diaz (2011), yellow right-pointing triangles from Kilian (1992), red big diamonds from Daflon et al. (1999, 2001a, 2001b), and white stars from Morel et al. (2008). The green line is the best fit found through an MCMC method, while the red lines represent the different fits also calculated with an MCMC algorithm. Only the objects in the sample were used for the fit; the Sun, Orion, 30 Dor, and I Zw 18 are also plotted as reference points.

(see Section 5.5). We find that only the C abundances determined from the equation with a negative slope match the trend suggested by the carbon abundances determined from the disk and halo stars (see Figure 11). We plotted in Figure 10 all of the linear fits with a negative slope (in red), as well as the best fit

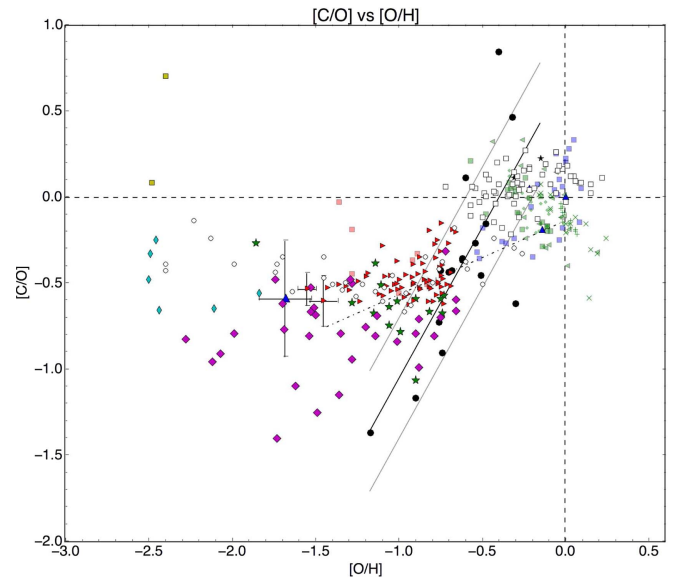


Figure 11. Behavior of $[C/O]$ vs. $[O/H]$ for various types of objects. Protostellar abundances were taken from Asplund et al. (2009). The filled circles are our STIS data, and the blue triangles are the same reference points as in Figure 7. Light red squares are the low-metallicity H II regions of Berg et al. (2016), and light blue squares are abundances obtained from the recombination lines of Esteban et al. (2002, 2004, 2014), García-Rojas & Esteban (2007), and López-Sánchez et al. (2007). Thin cyan diamonds are the C abundances of Pettini et al. (2008), yellow squares are very metal-poor DLAs of Cooke et al. (2015), and light green crosses, plus signs, dots, left triangles, and squares are disk stars of Cunha & Lambert (1994), Nieva & Simon-Diaz (2011), Kilian (1992), Daflon et al. (1999, 2001a, 2001b), and Morel et al. (2008), respectively. Smaller open circles are Galactic halo stars of Akerman et al. (2004) and white smaller squares are disk stars of Gustafsson et al. (1999). The following are approximated with Equation (12): green stars are extremely low-metallicity SDSS galaxies from James et al. (2015), red right-pointing triangles are the low-metallicity emission-line galaxy sample of Nava et al. (2006), and big magenta diamonds are very metal-poor DLAs of Pettini et al. (2008). The black line represents a linear fit to our data and the grey lines represent the errors of the fit; the black dotted-dashed line represents the linear fit of Garnett et al. (1995). For clarity, error bars are not presented, except for I Zw 18, for which the error bars represent the range of multiple results from the literature.

determined by the MCMC algorithm (in green). The equation for this line is

$$\log(C/N) = b - m \log(N/O), \quad (12)$$

where $b = 0.31 \pm_{0.12}^{0.15}$ and slope $m = 0.21 \pm_{0.11}^{0.09}$; the uncertainties were determined from the 25th and 75th percentiles. Of course, carbon abundances obtained with Equation (12) would be only a first crude approximation. To determine the uncertainty of this equation, we calculated the carbon abundances for our sample and compared these with the abundances obtained with the Garnett method. For the purpose of this analysis, we will call benchmark values the C abundances obtained with the Garnett method. We compared the approximated carbon abundances with the benchmark values and we obtained an average difference of 0.38 dex. Even though Equation (12) has a small statistical significance, it is interesting that the behavior of disk stars as well as the sample from Berg et al. (2016) also follow a negative slope. If this decrease of C/N with respect to N/O is true, further studies are needed to better characterize the behavior.

5.3. Abundance Discrepancy Factor

Abundances of photoionized objects are generally determined using CELs (see Section 1). However, in bright objects, oxygen and carbon abundances can also be determined with recombination lines (RLs). A well-known problem in the chemical analysis of photoionized objects is the discrepancy between the abundances determined with RLs and those determined with CELs (Peimbert et al. 1993, 2007; Esteban et al. 2004; García-Rojas & Esteban 2007; Peña-Guerrero et al. 2012a; Blanc et al. 2015, and references therein). This problem is generally referred to as the abundance discrepancy factor (ADF) problem, where ADF is defined as the ratio of abundances determined with RLs to those determined with CELs.

RLs yield higher abundances than CELs. Typical ADF values for H II regions lie in the 1.5 to 3 range (e.g., Peimbert et al. 1993; Peimbert 2003; Esteban et al. 2004; Nicholls et al. 2012; Peña-Guerrero et al. 2012a, 2012b) and in the 1.5 to 5 range (or higher than 20 in extreme cases) for most planetary nebulae (e.g., Liu & Danziger 1993; Nicholls et al. 2012; McNabb et al. 2013; Peimbert et al. 2014). There have been two major explanations for the ADF: (i) high-metallicity inclusions that will create cool high-density regions surrounded by hot low-density regions (e.g., Tsamis & Péquignot 2005), and (ii) thermal inhomogeneities in a chemically homogeneous medium that are caused by various physical processes such as shadowed regions, advancing ionization fronts, shock waves, magnetic reconnection, etc. (e.g., Peimbert & Peimbert 2011). A third explanation was recently proposed by Nicholls et al. (2012): electrons depart from a Maxwell-Boltzmann equilibrium energy distribution but can be described with a “ κ -distribution.” Nicholls et al. suggest that $\kappa \gtrsim 10$ is sufficient to encompass nearly all objects.

Our STIS abundances do not have the necessary resolution to accurately obtain abundances for carbon and oxygen via RLs. Esteban et al. (2004, 2014) determined the C and O abundances from RLs for a couple of objects in our sample. The comparison of the C and O abundances determined in this work with those determined by LSE09 and Esteban et al. (2014) are presented in Table 11, where column 1 is the galaxy name, column 2 the C abundances determined in this work with CELs, and column 3 the C abundances as determined in Esteban et al. (2014) with RLs. Column 4 shows the oxygen abundances determined in this work from CELs, column 5 the O abundance as determined in LSE09 also with CELs, and columns 6 and 7 present the O abundances as determined in Esteban et al. (2014) with CELs and RLs, respectively.

5.4. Dust Depletion

Depletion of heavy elements onto dust grains is important for the determination of accurate elemental abundances in the ISM (e.g., Garnett et al. 1995; Dwek 1998; Esteban et al. 1998; Peimbert & Peimbert 2010; Peña-Guerrero et al. 2012a). In the case of oxygen, depletion has been shown to be dependent on metallicity (Peimbert & Peimbert 2010): 0.09 ± 0.03 dex for $7.3 < 12 + \log(\text{O}/\text{H}) < 7.8$, 0.10 ± 0.03 dex for $7.8 < 12 + \log(\text{O}/\text{H}) < 8.3$, and 0.11 ± 0.03 for $8.3 < 12 + \log(\text{O}/\text{H}) < 8.8$. We have adopted this depletion correction for oxygen. However, the UV nature of the brightest carbon emission lines makes a dust depletion study particularly difficult. Studies of C depletion suggest a correction for the

nebular abundances from less than 0.1 (Sofia et al. 1994) to about 0.4 dex (Cardelli et al. 1993). Cunha & Lambert (1994) and Garnett et al. (1995) recommend a correction of 0.2 dex, independent of metallicity.

The dust-corrected carbon and oxygen abundances for our STIS sample are shown in Table 10. The correction due to dust depletion is almost about twice as high for carbon than for oxygen. We decided not to plot the corrected abundances since there is a possibility for both depletion corrections to be dependent on metallicity. If this is the case, the correction on oxygen would be more accurate than that for carbon. Nonetheless, the overall shape observed in Figure 7 is preserved when using corrected abundances, though values are slightly increased. The behavior of all other figures also follows this description: we find no significant change in the overall shape of the curves presented in this work. A possible consequence of the behavior of C/O versus O/H not being flat could be that the depletion of carbon has a metallicity dependence. The contribution of carbon from stars more massive than $25 M_{\odot}$ to the ISM strongly depends on metallicity: the higher the mass, the higher C is expelled into the ISM (Maeder 1992). However, the higher the metallicity of the gas, the greater the cross-section for dust radiation in the UV (Gustafsson et al. 1999), allowing for efficient destruction by photoionization. Hence, if there is a metallicity dependence in the C depletion, it is not a trivial one.

5.5. Damped Ly α systems

Damped Ly α systems (DLAs) are objects with high column density ($\log [N(\text{H I})] \geq 20.3 \text{ cm}^{-2}$) of predominantly neutral gas detected in the spectra of an unrelated background light source, typically a quasar (Cooke et al. 2015). DLAs have acquired particular attention mainly because (i) the most metal-poor DLAs offer a unique opportunity to study the enrichment of galaxies due to the first generations of stars (Kobayashi et al. 2011), and (ii) DLAs appear to sample various types of galaxies, from those with an extended HI disk to subgalactic-size halos (Wolfe et al. 2005) at a wide range of redshifts.

The dominant neutral gas component of DLAs allows for the measurement of heavy-element abundances to be straightforward, without the need for large ionization corrections. However, chemical abundances of carbon, nitrogen, and oxygen have received little attention in comparison to other heavy elements such as Cr, Fe, Mg, or Zn, though the relative abundances of C, N, and O, particularly at low metallicities, provide extremely valuable information about early nucleosynthesis stages (Wolfe et al. 2005; Cooke et al. 2015, and references therein). Since C and O are abundant elements with strong atomic transitions, their corresponding absorption lines are strongly saturated, thus making them unusable for abundance determination (Pettini et al. 2008). The N absorption lines tend to be weak and blended with intergalactic Ly α (Ly α) forest lines (Pettini et al. 1995, 2002). Nonetheless, the low metallicity and simple velocity structure of DLAs facilitate the measurement of C, N, and O abundances (Pettini et al. 2008).

Several previous studies (e.g., Dessauges-Zavadsky et al. 2003; Péroux et al. 2007; Pettini et al. 2008; Cooke et al. 2011, 2015) have obtained C/O measurements from unsaturated C II and O I absorption lines. Furthermore, it has recently been suggested that DLAs have chemical evolution and kinematic structure that resemble those of Local Group dwarf galaxies (Cooke et al. 2015). We included all previously cited carbon

Table 11
Comparison of Total C and O Abundances with Literature Values^a

Galaxy Name	12+log(C/H)		12+log(O/H)			
	This Work ^b	Est14 ^c	This Work ^b	Lop09 ^d	Est14 ^c	
	CELs	RLs	CELs	CELs	CELs	RLs
Mrk 960	7.49 ± 0.20	...	8.11 ± 0.19	8.37 ± 0.10
SBS 0218+003	6.82 ± 0.19	...	7.99 ± 0.06	7.95 ± 0.05
Mrk 1087	8.60 ± 0.19	...	8.41 ± 0.16	8.57 ± 0.10
NGC 1741	7.83 ± 0.20	<7.97 ^e	8.25 ± 0.10	8.22 ± 0.05	8.41 ± 0.06 ^c	<8.68 ^e
Mrk 5	6.97 ± 0.19	...	7.96 ± 0.10	8.07 ± 0.07
Mrk 1199	8.91 ± 0.23	...	8.33 ± 0.22	8.57 ± 0.21
IRAS 08208+2816	7.54 ± 0.15	...	8.43 ± 0.11	8.41 ± 0.10
IRAS 08339+6517	8.27 ± 0.19	...	8.42 ± 0.07	8.45 ± 0.10
SBS 0926+606A	7.29 ± 0.15	...	7.98 ± 0.09	7.94 ± 0.08
Arp 252	7.98 ± 0.19	...	8.13 ± 0.15	8.46 ± 0.11
SBS 0948+532	7.33 ± 0.15	...	8.03 ± 0.09	8.03 ± 0.05
Tol 9	8.54 ± 0.21	...	8.58 ± 0.15	8.58 ± 0.15
SBS 1054+365	7.48 ± 0.11	...	8.11 ± 0.09	8.00 ± 0.07
POX 4	7.36 ± 0.08	8.11:	8.05 ± 0.05	8.03 ± 0.04	8.09 ± 0.02	8.24:
SBS 1319+579	7.66 ± 0.17	...	8.19 ± 0.09	8.05 ± 0.06
SBS 1415+437	5.92 ± 0.21	5.42 ± 0.03 ^f	7.56 ± 0.07	7.61 ± 0.06
Tol 1457-262	6.40 ± 0.30	...	7.83 ± 0.16	8.05 ± 0.07	8.21 ± 0.02	...
III Zw 107	7.50 ± 0.18	...	8.22 ± 0.10	8.23 ± 0.09

Notes.

^a Comparison of the total C and O abundances derived in this work using the direct method with those previously reported in the literature. Abundances are presented in units of 12+log(X/H), and considering a homogeneous thermal structure ($t^2 = 0.000$).

^b Abundances determined with the method described in Garnett et al. (1995).

^c Abundances taken from Esteban et al. (2014). The oxygen abundances presented from RLs are those values for O²⁺.

^d Abundances taken from LSE09.

^e Abundances taken from Esteban et al. (2004) from the optical RL ionic abundance of X²⁺.

^f Abundance taken from James et al. (2014) from the C II UV absorption lines for the neutral ISM. The neutral ISM value is lower than that from the H II regions by ~0.5 dex, which agrees with the results of Aloisi et al. (2003) and Leboutteiller et al. (2009).

measurements in the [C/O] versus [O/H] diagram in Figure 11. In addition, we also included in this figure the sample of James et al. (2015), which is a subsample of 12 extremely metal-poor galaxies morphologically selected from the SDSS, as well as the sample from Nava et al. (2006), which is a compiled sample of low-metallicity emission-line galaxies. We calculated a crude first approximation of the C/O values for these two literature samples from their N/O and O/H values and Equation (12). Even though the scatter is large in Figure 11, the overall shape of the figure with the resulting carbon abundances from Equation (12) seem to agree with similar figures in the literature, e.g., Figure 7 of Berg et al. (2016). Though the uncertainties are large, if we take the center values to be true, it becomes apparent that different types of objects “prefer” certain areas of the [C/O] versus [O/H] diagram. For the carbon abundances we gathered from the literature and the ones we derived in this work (either with the Garnett method or with our linear approximation), we observe that the most metal-poor DLAs are in the region $-0.7 < [C/O] < 0.7$ and $[O/H] < -1.8$, higher metallicity DLAs are in the region $-1.5 < [C/O] < 0.0$ and $-2.5 < [O/H] < -0.5$, extremely low-metallicity galaxies are in the region $-1.0 < [C/O] < 0.0$ and $-2.0 < [O/H] < -0.5$, low-metallicity emission-line galaxies are in the region $-0.5 < [C/O] < 0.0$ and $-1.5 < [O/H] < -0.5$, neutral ISM measurements with the 0.5 dex addition to make the values comparable to those from star-forming regions (James et al. 2014) are in the region $-2.0 < [C/O] < -1.0$ and $-2.0 <$

$[O/H] < -0.5$, halo stars are in the region $-1.0 < [C/O] < 0.0$ and $-2.5 < [O/H] < -0.5$, disk stars are in the region $-0.4 < [C/O] < 0.4$ and $-0.7 < [O/H] < 0.3$, and W-R galaxies are in the region $-1.5 < [C/O] < 1.0$ and $-1.3 < [O/H] < 0.0$. The translation of Equation (9) into solar values is the following:

$$[C/O] = m [O/H] + b, \quad (13)$$

where $m = 1.75 \pm 0.35$ and $b = 0.69 \pm 0.22$. We also translated the linear fit of Garnett et al. (1995) to solar values, and we present it as well in Figure 11. Metal abundances in the neutral ISM can be determined with far-UV absorption lines, which requires using bright UV sources as background spectra whose light is absorbed along the line of sight, similar to the study of DLAs (Lu et al. 1996). To compare our results with the C determinations from absorption lines, it is important to consider that the analysis of the *Far Ultraviolet Spectroscopic Explorer (FUSE)* spectra of I Zw 18 with this technique indicates that the abundances of the alpha elements such as O, Ar, Si, and N are ~0.5 dex lower in the neutral ISM than in the H II regions, while the abundance of Fe remains the same (Aloisi et al. 2003), which has been confirmed by several other studies (e.g., Leboutteiller et al. 2009 and references therein).

Previous and current studies have reported large carbon enhancements in DLAs with C/O values matching those of halo stars of similar metallicity or even higher values, which is

not expected from Galactic chemical evolution models based on conventional stellar yields (e.g., Pettini et al. 2008; Esteban et al. 2014; Cooke et al. 2015, 2015, and references therein). Such carbon enhancements suggest higher stellar carbon yields, probably originating from the stellar rotation, which promotes mixing in the stellar interiors (Pettini et al. 2008). This could also be taken as independent confirmation of the non-flat behavior of C/O with respect to O/H as explained by metallicity-dependent stellar yields (Garnett et al. 1995). Moreover, Akerman et al. (2004) suggested that [C/O] values could not remain constant at [C/O] = -0.5 , as previously thought, but increase again to approach solar metallicities at about [O/H] ~ -3 , which would be due to metallicity-dependent non-LTE corrections to the [C/O] ratio. They proposed Population III stars as a possible explanation for the near-solar values of [C/O] at low metallicities, particularly if assuming a top-heavy IMF for these stars. Akerman et al. suggested that the higher temperatures reached in the cores of metal-free stars could shift the balance in the carbon and oxygen reactions, consequently producing a higher carbon yield, or that the mixing and fallback models of high-energy supernova explosions of Umeda & Nomoto (2002, 2005) could be responsible for the carbon enhancement at early times. In either case, Figure 11 confirms the behavior predicted by Akerman et al. (2004) for their “standard” model, which uses the yields of Meynet & Maeder (2002) for massive stars and those of van den Hoek & Groenewegen (1997) for low- and intermediate-mass stars, combined with the metal-free yields of Chieffi & Limongi (2002). Furthermore, Figure 11 also supports the behavior of [C/O] versus [O/H] observed by Bensby & Feltzing (2006), Pettini et al. (2008), Esteban et al. (2014), and Berg et al. (2016): Bensby & Feltzing (2006) suggest that the higher values of C/O at higher metallicities could be due in the most part by low and intermediate-mass stars; Pettini et al. (2008) pointed out that the increase of the C/O ratio at lower metallicities suggests an additional source of carbon from the massive stars responsible for early nucleosynthesis; Esteban et al. (2014) explain that the positions of star-forming dwarf galaxies coinciding with those of Galactic halo stars suggests the same origin for the bulk of carbon in those galaxies; and Berg et al. (2016) argue that variations in the IMF could contribute to the large dispersion in the C/O values.

The characteristic H I Ly α absorption lines observed in DLAs are broadened by radiation damping, yet in some objects, emission lines can also be observed. Of these emission lines, Ly α (1216 Å) is the most valuable spectroscopic star-forming indicator in the redshift range of $4 < z < 6$ (Stark et al. 2011). At higher redshifts, this line is no longer a reliable star-forming indicator due to the resonant scattering by neutral gas in the IGM (Zitrin et al. 2015). The galaxy population at $z > 6$ has in general lower UV luminosities and stellar masses than those from samples at $z \simeq 2-3$, as well as large star formation rates, indicating a rapidly growing young stellar population (e.g., Stark et al. 2015b and references therein). Among the strongest emission lines of early galaxies are [O III] 5007 Å and H α ; however, at $z \sim 6$, these lines are situated at about 3–5 μm , which makes them non-detectable with ground-based telescopes. Nonetheless, other UV emission lines such as [O III] 1660+6 and [C III] 1907+09 can probe the ionizing spectrum of galaxies at $z \gtrsim 7$. Stark et al. (2014)

reported tentative detections of [C III] in two galaxies with z of 6.029 and 7.213 (from Ly α), while Zitrin et al. (2015) reached a 5σ median flux limit for [C III] for an integration of 5 hours in the H -band in their pilot survey of the reionization era. This suggests that in the near future, points with $z > 6$ will be added to the C/O versus O/H diagram, giving us a more extended view of the carbon enrichment of the Universe.

5.6. MCMC Modeling

Due to the low resolution of our STIS observations, we wanted an independent way to determine if our carbon measured through the Garnett method and the oxygen abundances measured through the direct method were sensible. We used the MCMC technique to explore the parameter space and see where our carbon and oxygen abundances lie with respect to several thousand independent modeled samples with similar physical conditions.

We used Cloudy for the photoionization models and an ionization spectrum input from Starburst99. The chain ran using the emcee algorithm. We obtained about 30,000 photoionization models per object. We calculated χ^2 from comparing the observed line intensities to the modeled ones. We used the intensity relative to H β of nine lines to determine a χ^2 value per model: H I 4340, 4861, and 6563; He I 5876; He II 4686; [O II] 3727; [O III] 5007; [C III] 1909; and [S III] 9532 Å. To account for the observed oxygen temperatures, we made a subsample of models, taking only those with temperatures of [O III] and [O II] $T_{\text{model}} \leq T_{\text{observed}} \pm 2500$ K, and we took the average of this subsample. The uncertainties for the average final abundances were obtained from the 25th and 75th percentiles. For the modeling setup and specific code versions used, please see the Appendix.

The bulk of the final average abundances obtained from the models for carbon and oxygen agree with our measured abundances, within the measurements’ errors. Due to using almost no constraints for the runs, we obtained large values for the 25th and 75th percentiles. Nonetheless, we noticed that the abundances of nitrogen, neon, and sulphur are not in close agreement with the observations. This is due to the loose constraints we used for the models, and since our STIS sample is composed of objects with a top-heavy IMF, it is a potential hint that the IMF we used for the models may not be the most adequate one. This issue requires further study to determine if it is true.

We combined all models from all objects in the sample to create a stacked log(C/O) versus $12 + \log(\text{O}/\text{H})$ diagram, and a log(N/O) versus log(C/N) diagram. We noted that the main value of the models is between -0.9 and -0.8 , which is interestingly coincidental with [C/O] for metal-poor halo stars according to Tomkin et al. (1992; or -0.9 and -0.8 adopting the protosolar values for $\text{C}/\text{O}_{\odot} = -0.26 \pm 0.07$ and $12 + \log(\text{O}/\text{H})_{\odot} = 8.73 \pm 0.05$ from Asplund et al. 2009), and with the carbon abundances determined via the RLs of Esteban et al. (2014). If the IMF does indeed play an important role in the C/O versus O/H diagram, then this result would imply that there is a specific behavior for objects that have a similar IMF, and that halo stars are well described by a Kroupa IMF, like the one we used for our models.

6. Summary and Conclusions

We obtained STIS spectra covering the spectral region from about 1600 to 10000 Å for 18 starburst galaxies selected from the sample of WR galaxies discussed by LSE08, LSE09, LSE10a, and LSE10b. Our goal is to study the enhancement of carbon in the ISM due to massive stars. We obtained physical conditions and chemical abundances for these 18 objects through standard nebular analysis. To determine the carbon abundances, we used the method described in Garnett et al. (1995). The main results of the present work are as follows.

1. We confirm our previous results: there is an increase in C/O with respect to O/H, yet we do not find a simple correlation. The most likely explanation for the non-constant relation (predicted as constant at low O/H by instantaneous recycling models for both carbon and oxygen) is that the yield of C varies with respect to O. Furthermore, our results indicate that the nucleosynthesis of carbon and/or oxygen deviates from the closed-box model, at least when dealing with objects with a clear “starprint” of massive stars (i.e., W-R stars). This behavior agrees with the results of LSE10c, who also found that their galaxy sample did not agree with the closed-box model. These authors argue that the pristine gas inflow or the enriched gas outflow played an important role in the chemical evolution of their sample galaxies. When comparing our STIS sample C/O measurements with other references in the literature, such Berg et al. (2016), we find that there is a steeper slope of C/O versus O/H for our data, suggesting that the top-heavy IMF might have an effect on the carbon production, i.e., when massive stars are numerous, there is an additional contribution of C into the ISM for objects with metallicities higher than $12+\log(\text{O}/\text{H}) \gtrsim 8.0$.
2. Our data suggest that the N/C ratio increases with increasing carbon abundance. This contradicts the behavior of the sample presented in Berg et al. (2016), again suggesting that the IMF has a strong influence on the carbon production: at metallicities higher than $12+\log(\text{O}/\text{H}) \gtrsim 8.0$, massive stars contribute more to the production of C. Further data are required to characterize this correlation, if it indeed exists.
3. We find a potential empirical correlation between $\log(\text{C}/\text{N})$ with respect to $\log(\text{N}/\text{O})$. This relation estimates the carbon abundance from measurements of oxygen and nitrogen abundances, but should only be taken as a first-order approximation. The average difference in the carbon abundances approximated for our sample with this equation with respect to the carbon abundances obtained with the Garnett method for the same sample is 0.38 dex.
4. In this work, we used an MCMC method to determine whether our carbon and oxygen abundance measurements were sensible. This method permits the parameter space to be explored. However, to obtain accurate results with this technique, detailed photoionization models are required, but they can provide an effective and efficient technique to study correlations and/or degeneracies between abundances within an object as shown by Tremonti et al. (2004), Pérez-Montero (2014), and Blanc et al. (2015).
5. The average value of $\log(\text{C}/\text{O})$ from all Cloudy models is about -0.8 , which coincides very well with the main value of $\log(\text{C}/\text{O})$ for halo stars. If the IMF indeed has a strong effect on the production of carbon, the behavior shown by the Cloudy models indicates that the IMF of the models “promoted” a greater number of intermediate-mass stars rather than massive stars; hence, the nucleosyntheses of carbon and nitrogen are most likely due to the same stars. The coincidence of C/O values could be an indication that halo stars are well described by a Kroupa IMF.
6. The addition of DLAs, disk and halo stars, and neutral ISM to the [C/O] versus [O/H] diagram provides additional insight into the carbon enrichment of the universe with respect to oxygen. Independent results from different types of objects may confirm that the observed trends are due to stellar yields being metallicity dependent rather than the instantaneous recycling assumption not holding true.
7. From the carbon determinations we compiled from the literature and those we determined in this work, we observe that different types of objects seem to be located in specific regions of the [C/O] versus [O/H] diagram. This diagram confirms the suggested behavior of [C/O] at lower metallicities observed by Pettini et al. (2008) and Esteban et al. (2014), and predicted by Akerman et al. (2004), which is likely due to Population III stars, before nucleosynthesis from Population II takes over, and agrees with Berg et al. (2016) that the scatter of the C/O values are likely due to differences in the IMFs.

Our results indicate that carbon and/or oxygen nucleosynthesis deviates from the instantaneous recycling and closed-box models, at least in the presence of a large number of massive stars. The difference in the steep slope we find in the behavior of $\log(\text{C}/\text{O})$ with respect to $12+\log(\text{O}/\text{H})$ versus previous studies of C/O in other objects suggests that carbon production is indeed greatly affected by the presence of massive stars. The behavior of C/O with metallicity resembles the relation between a primary and a secondary element, where the abundance ratio of the secondary to the primary element is predicted to increase with the abundance of its seed. A classic example of such a behavior is that for the N/O to O/H ratio; see, e.g., Nava et al. (2006). The most plausible explanation for this behavior between C/O and O/H is that carbon is returned to the ISM by intermediate-mass stars on longer timescales compared to oxygen, which is mainly returned to the ISM by massive stars; hence, C/O increases as O/H increases. This effect is amplified by the metallicity dependence of the carbon yields. Nonetheless, our measurements indicate that intermediate-mass stars play a dominant role in the production of carbon in the range of $-2.5 \lesssim [\text{O}/\text{H}] \lesssim -0.5$.

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Appendix MCMC Modeling of Photoionized Objects

To verify and further understand the behavior of the chemical abundances in our sample, we ran a series of `Cloudy` models with a `Starburst99` synthetic spectrum as the ionization source through an MCMC method (see Appendix A.0.1). We ran about 30,000 models per object, letting the chain freely explore the parameter space, with the purpose of not creating ad hoc models for our observations. We then restricted T_e in order to select those models with closer physical conditions and averaged this subsample to obtain our best estimate for the modeled abundances. The bulk of these values agreed with the abundances of O and C derived from our observations within their uncertainties. Other works, such as Brinchmann et al. (2008), Tremonti et al. (2004), Pérez-Montero (2014) and Blanc et al. (2015), also used a similar Bayesian analysis to determine the physical conditions and chemical composition of star-forming galaxies. Nonetheless, we have chosen to determine our best estimate of the chemical abundances and present our results in a different way.

A.0.1. Setup for the MCMC Photoionization Models

The main goal of the analysis of a problem with an MCMC method is to efficiently sample a set of parameters, using random numbers drawn from a uniform probability in a given range. As a result, probability density functions (PDFs) are obtained for a set of parameters. Such PDFs then show a correlation, anti-correlation, or no relation at all between the parameters. Essentially, the MCMC itself consists of random walks that provide a sample of the posterior probability distribution, given certain priors and according to Bayes' law: the posterior probability is proportional to the product of the likelihood and the prior probability,

$$P(\theta, T_e(\text{O}), n_e|D) \propto \mathcal{L}(\theta, T_e(\text{O}), n_e) P(D|\theta, T_e(\text{O}), n_e), \quad (14)$$

where θ represents the set of abundance ratios of He/H, O/H, C/O, N/O, Ne/O, and S/O; $T_e(\text{O})$ represents the electron temperatures of [O II] and [O III]; n_e represents the electron density; and D represents the possible true values.

To determine the chemical composition of a set of objects, we used 100 walkers and conducted 100 runs per object, which resulted in about 30,000 models of `Cloudy` version 13.03 (Ferland et al. 2013) per object. We used the Python module `pyCloudy` version 0.8.37 (Morisset 2013) to include `Cloudy` into our code. The priors were set so that the trial abundances (or random walks) were physically plausible values for the set of abundances (or dimensions). The ionizing source we used for the `Cloudy` models was a synthetic spectrum generated with `Starburst99` version 7.0.1 (Leitherer et al. 2014) at an age of 10^8 years. We used a continuous SFR for the `Starburst99` spectra following Goldader et al. (1997), who find that an instantaneous SF burst provides an unrealistic short range of ages for their sample of luminous infrared galaxies, and Kewley et al. (2001), who find that a continuous SFR agrees better with their sample of infrared starburst galaxies, the spectra of which are dominated by emission from

W-R stars. Additional conditions used for the `Starburst99` spectra were a Kroupa IMF with two exponents of 1.3 and 2.3 (Kroupa 2014) and mass boundaries of 0.1, 0.5, and $120 M_\odot$; the 2012/13 Geneva stellar evolutionary tracks with rotation velocities of 40% of the break-up velocity; and $Z = 0.014$ (Ekstrom et al. 2012; Georgy et al. 2013). All other parameters were set to the default values in `Starburst99` (see <http://www.stsci.edu/science/starburst99/docs/code.html>).

A.0.2. MCMC Photoionization Models Input and Comparison to Observations

The input trial abundances for the `Cloudy` models were varied according to the MCMC algorithm in the Python module `emcee` (Foreman-Mackey et al. 2013). The `Cloudy` output contains line intensities as well as temperatures of [O III] and [O II]. We assumed that the uncertainties of the observations follow a Gaussian distribution; therefore, we used χ^2 as a likelihood function—more specifically, we used $\ln(-\chi^2/2)$.

We calculated χ^2 from comparing the observed line intensities to the modeled ones. We used the intensity relative to H β of nine lines to determine a χ^2 value per model: H I 4340, 4861, and 6563; He I 5876; He II 4686; [O II] 3727; [O III] 5007; C III] 1909; and [S III] 9532 Å. For the `Cloudy` models, we used the following options: electron density of 150 cm^{-3} , cosmic-ray background, atom H-like levels, atom He-like levels, no fine opacities, no molecules, and no level 2 lines.

An initial chemical composition was also given as input for each model. One of the advantages of the MCMC method is that the initial guess does not have an effect on the “final” answer because the set of walkers obtained at the end should be independent samples from the distribution (for a more detailed discussion on the initial guess, see Foreman-Mackey et al. 2013). We initiated the chains with random abundance values in the following ranges of typical observed values for H II regions: $9.8 \leq \text{He} \leq 10.0$, $8.3 \leq \text{O} \leq 8.6$, $-0.3 \leq \text{C/O} \leq 0.1$, $-1.3 \leq \text{N/O} \leq -1.0$, $-0.3 \leq \text{Ne/O} \leq 0.0$, and $-2.0 \leq \text{S/O} \leq -1.7$, in units of $12+\log(\text{X}/\text{H})$ and $\log(\text{X}/\text{O})$, respectively. The initial abundances were varied according to the MCMC algorithm. These were checked against the prior probability (allowed values): $9.5 \leq \text{He} \leq 12.0$, $7.5 \leq \text{O} \leq 8.7$, $-1.6 \leq \text{C/O} \leq 1.7$, $-1.7 \leq \text{N/O} \leq -0.4$, $-1.0 \leq \text{Ne/O} \leq 0.01$, and $-2.3 \leq \text{S/O} \leq -1.4$. We used a top-hat function for each of the abundances:

$$f(x) = \begin{cases} -\ln(b - a) & a < x < b \\ -\infty & \text{else} \end{cases}, \quad (15)$$

where a and b are the abundance values corresponding to the allowed minimum and maximum, respectively. If the state is accepted, then the values are converted to abundances relative to hydrogen in order to be sent as input for the `Cloudy` models. With the mentioned options, each model took between 3 and 5 minutes to run.

A.1. Results from the MCMC Photoionization Modeling

The result of the MCMC is a distribution of probabilities of occurrence for a set of parameters. Figure 12 (and the accompanying figure set) shows the relations between our parameters: He/H, O/H, C/O, N/O, Ne/O, and S/O, in units of $12+\log(\text{X}/\text{H})$ or $\log(\text{X}/\text{O})$. The blue crosses represent the

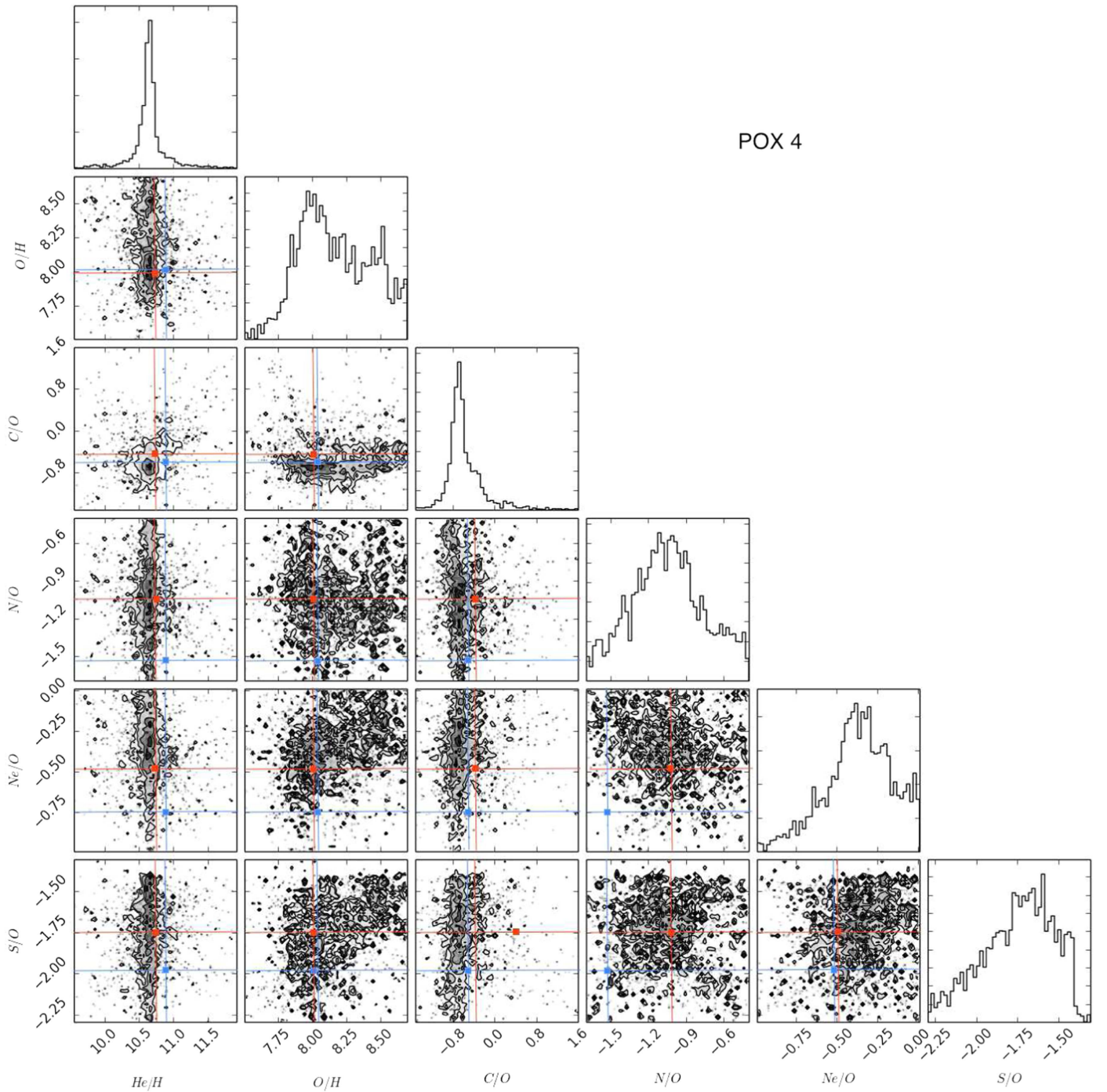


Figure 12. Behavior of the probability of occurrences for the set of abundances of $12+\log(\text{He}/\text{H})$, $12+\log(\text{O}/\text{H})$, $\log(\text{C}/\text{O})$, $\log(\text{N}/\text{O})$, $\log(\text{Ne}/\text{O})$, and $\log(\text{S}/\text{O})$ according to our MCMC photoionization modeling for POX 4. The blue crosshairs represent the measurements from the observations, and the red crosses represent our best estimate from the models. The uppermost panel in each column shows the histogram of the variable directly below it. The figure was created with `corner.py` (Foreman-Mackey 2016).

(The complete figure set (18 images) is available.)

benchmark abundances (i.e., abundances obtained from our STIS observations), and the red crosses represent the best estimate from the models. There are various methods to obtain a value from this distribution in order to compare it with the measured abundances from observations. One such method is to maximize the likelihood and use that value (particularly useful when the histogram presents a Gaussian distribution; see

Figure 12). However, the major advantage of the MCMC analysis is the availability of a distribution of values. The average or median values use such a distribution to obtain a best estimate. In this work, we used the average values; the median values yield essentially the same numbers.

In order to account for the observed oxygen temperatures, we made a subsample of models taking only those with

Table 12
Chemical Composition of the Modeled Sample^a

Galaxy Name	He/H	O/H	C/H	N/H	Ne/H	S/H
Mrk 960	10.89 \pm _{0.34} ^{0.40}	8.50 \pm _{0.07} ^{0.14}	8.29 \pm _{0.60} ^{0.64}	7.34 \pm _{0.20} ^{0.11}	8.04 \pm 0.21	6.78 \pm _{0.17} ^{0.20}
SBS 0218+003	10.81 \pm 0.05	7.86 \pm _{0.11} ^{0.09}	6.90 \pm _{0.27} ^{0.11}	6.79 \pm 0.14	7.46 \pm 0.14	6.09 \pm 0.18
Mrk 1087	10.65 \pm _{0.39} ^{0.64}	8.46 \pm _{0.08} ^{0.12}	8.84 \pm _{0.61} ^{0.65}	7.37 \pm _{0.29} ^{0.23}	7.97 \pm 0.23	6.71 \pm 0.21
NGC 1741	10.99 \pm 0.48	8.05 \pm 0.17	7.60 \pm _{0.37} ^{0.34}	6.99 \pm _{0.22} ^{0.18}	7.58 \pm 0.20	6.27 \pm 0.13
Mrk 5	10.70 \pm _{0.06} ^{0.03}	8.09 \pm _{0.07} ^{0.09}	7.62 \pm _{0.27} ^{0.20}	7.10 \pm 0.14	7.64 \pm _{0.09} ^{0.12}	6.31 \pm 0.12
Mrk 1199	11.19 \pm _{0.53} ^{0.52}	8.59 \pm _{0.05} ^{0.08}	9.07 \pm _{0.27} ^{0.38}	7.62 \pm _{0.28} ^{0.14}	8.48 \pm _{0.08} ^{0.11}	7.03 \pm 0.15
IRAS 08208+2816	10.86 \pm 0.11	8.51 \pm _{0.07} ^{0.10}	7.51 \pm _{0.09} ^{0.09}	7.66 \pm _{0.20} ^{0.23}	8.19 \pm _{0.15} ^{0.17}	6.69 \pm _{0.18} ^{0.22}
IRAS 08339+6517	11.28 \pm _{0.29} ^{0.37}	8.52 \pm _{0.10} ^{0.23}	8.85 \pm _{0.54} ^{0.60}	7.44 \pm 0.28	8.23 \pm _{0.09} ^{0.17}	6.79 \pm 0.16
SBS 0926+606A	10.86 \pm _{0.09} ^{0.19}	7.90 \pm 0.09	7.38 \pm _{0.11} ^{0.03}	6.81 \pm 0.18	7.44 \pm 0.14	6.12 \pm 0.10
Arp 252	10.62 \pm _{0.36} ^{0.43}	8.42 \pm _{0.15} ^{0.17}	8.81 \pm _{0.41} ^{0.60}	7.50 \pm _{0.14} ^{0.21}	8.08 \pm _{0.23} ^{0.19}	6.64 \pm 0.11
SBS 0948+532	10.58 \pm _{0.13} ^{0.07}	8.06 \pm 0.14	7.28 \pm _{0.23} ^{0.12}	6.94 \pm 0.22	7.65 \pm _{0.16} ^{0.18}	6.24 \pm 0.18
Tol 9	11.26 \pm _{0.51} ^{0.45}	8.56 \pm _{0.07} ^{0.10}	8.44 \pm _{0.42} ^{0.58}	7.57 \pm _{0.21} ^{0.18}	8.36 \pm _{0.18} ^{0.18}	6.68 \pm 0.26
SBS 1054+365	10.83 \pm _{0.34} ^{0.40}	8.03 \pm 0.13	7.46 \pm _{0.33} ^{0.43}	7.00 \pm _{0.18} ^{0.20}	7.59 \pm 0.14	6.23 \pm _{0.12} ^{0.14}
POX 4	10.70 \pm _{0.14} ^{0.10}	8.01 \pm 0.14	7.51 \pm _{0.25} ^{0.15}	6.92 \pm _{0.19} ^{0.17}	7.50 \pm 0.12	6.26 \pm 0.13
SBS 1319+579	10.66 \pm 0.07	8.09 \pm _{0.11} ^{0.14}	7.46 \pm _{0.18} ^{0.11}	7.11 \pm 0.15	7.56 \pm 0.12	6.25 \pm 0.14
SBS 1415+437	10.70 \pm 0.23	7.73 \pm _{0.13} ^{0.11}	6.44 \pm _{0.20} ^{0.05}	6.50 \pm _{0.17} ^{0.15}	7.29 \pm _{0.20} ^{0.22}	5.82 \pm _{0.12} ^{0.10}
Tol 1457–262	11.31 \pm _{0.01} ^{0.18}	7.92 \pm _{0.05} ^{0.07}	6.99 \pm _{0.11} ^{0.06}	6.82 \pm 0.19	7.61 \pm _{0.09} ^{0.12}	6.02 \pm 0.13
III Zw 107	10.91 \pm _{0.19} ^{0.28}	8.25 \pm 0.11	7.87 \pm 0.33	7.21 \pm 0.17	7.96 \pm 0.17	6.44 \pm 0.16

Note.

^a In units of $12+\log n(X)/n(H)$.

temperatures of [O III] and [O II] $T_{\text{model}} \leq T_{\text{observed}} \pm 2500$ K, and we took the average of this subsample. We chose an arbitrarily large range of temperatures in order to encompass all models within the range of uncertainties of the observations. Table 12 presents the gaseous abundances in units of $12+\log(X/H)$ or $\log(X/O)$ of He/H, O/H, C/H, N/H, Ne/H, and S/H as obtained from the MCMC Photoionization Modeling method for the same elements. The uncertainties we present for the modeled abundances were obtained from the 25th and 75th percentiles of the subsample of models (i.e., all models with $T_{\text{observed}} \pm 2500$ K). These uncertainties are large, due to our decision of letting the chain explore the parameter space as freely as possible. For reference, we created four spectral zoom-in windows that show the major emission lines per object. Figure 13 is an example of these zoom-in windows (the corresponding figures for all other objects are presented in the figure set).

A.2. Analysis of the MCMC Photoionization Models

The C/O ratios obtained from the MCMC photoionization models are plotted against the gaseous modeled oxygen abundance in Figure 14. We see a similar behavior to that shown in Figure 7. The uncertainties are large due to the freedom with which the MCMC was run. We did this with the purpose of exploring most of the parameter space. Figure 12 shows an example of the PDF relations within the set of parameters for each object. For POX 4, we find that there are three instances with no correlation: (i) the panel of C/O versus He, which shows a circular contour, (ii) panels such as S/O versus Ne/O, which show points in the entire space, and (iii) panels such as S/O versus C/O, in which one of the parameters has a constant value. On the other hand, panels like Ne/O versus O show a potential linear correlation. In this figure, we see no anti-correlations.

From our MCMC models (i.e., the PDF diagrams like Figure 12), we observed that for the $\log(C/O)$ versus $12+\log(O/H)$ diagram, about half of the objects in our sample appear to show an increasing slope whereas the other half appear to show a somewhat flat behavior. We determined the slopes of this diagram for the models ran per galaxy and found that the range of the slopes is -0.6 to 1.0 , with an average slope of 0.4 and a median slope of 0.5 . Figure 15 shows the slopes from our MCMC photoionization models per galaxy. Each slope was determined from a linear fit to all models for that object in the $\log(C/O)$ versus $12+\log(O/H)$ diagram. We observe that these slopes do not seem to be related to metallicity, distance, redshift, carbon abundance, type of object, interaction degree, S/N of the pertinent lines, number of W-R stars versus O stars, or ionization degree. To check if this was a random effect of the MCMC, we re-ran the chain three times with the same conditions for three objects (the least and most metallic objects, and an intermediate-metallicity object: SBS 1415+437, Tol 9, and SBS 0218+003, respectively). The behavior of the slope remained the same for all three objects. We then decided to combine all of the models for all of the objects into a single figure of C/O versus $12+\log(O/H)$. The result is presented in Figure 16, which shows a clear flat behavior with the main or most probable value of C/O of about -0.8 . This flat behavior in the context of the MCMC analysis suggests that there is no obvious correlation between $\log(C/O)$ and $12+\log(O/H)$. Nonetheless, the value of C/O seems to be well restricted to values between -1.6 and 0.0 . The bulk of the sample in this paper has C/O values within this range. From the analysis of Figures 7–9, there seems to be a specific commonality that separates points from our sample from literature references, and we suspect that could be the top-heavy IMF that all objects in our sample share. If this is true, the slope behavior we observe in the models is to be expected since the IMF was the same for

POX 4

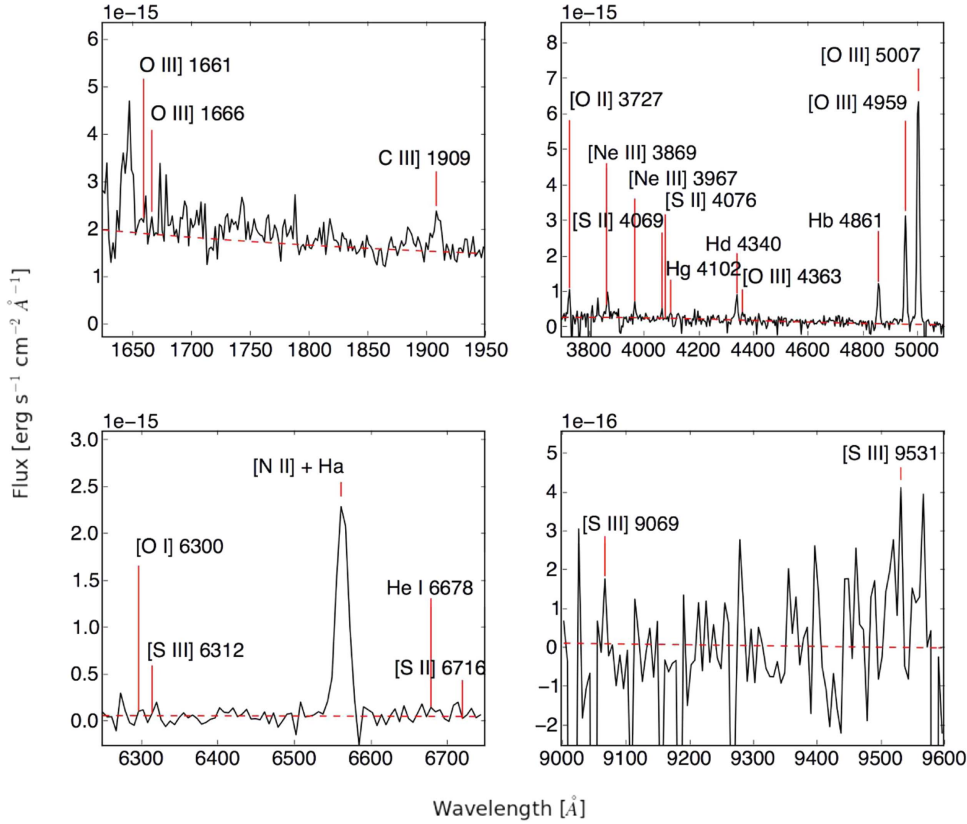


Figure 13. Zoom-in of the spectrum for the UV, blue, red, and IR regions for POX 4. (The complete figure set (18 images) is available.)

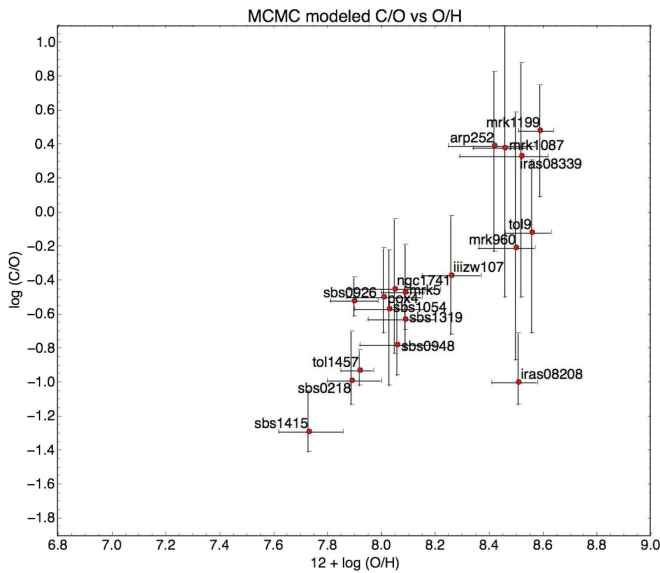


Figure 14. Behavior of C/O vs. $12+\log(\text{O}/\text{H})$ as obtained from MCMC. Symbols are the same as in Figure 7.

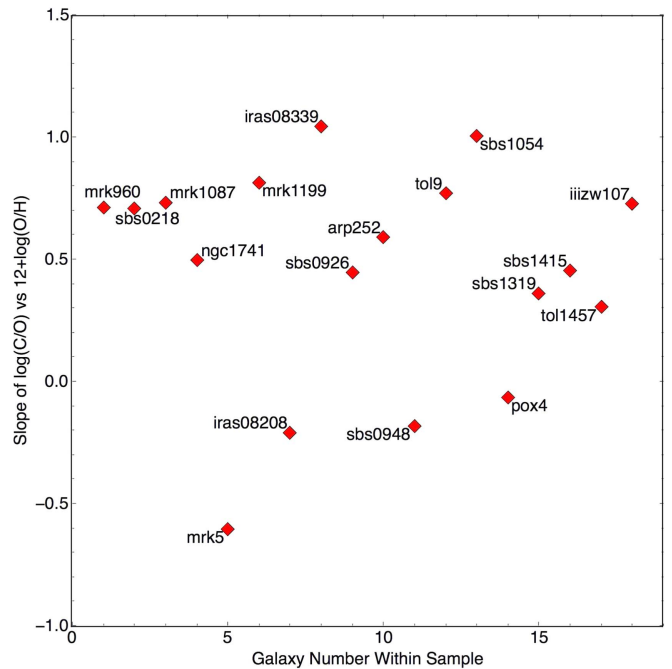


Figure 15. Behavior of the slopes of $\log(\text{C}/\text{O})$ to $12+\log(\text{O}/\text{H})$ as obtained from the MCMC photoionization modeling per galaxy.

all models. A separate study would be required to assess if this is indeed the case.

Using all of the models for all of the objects, we also studied $\log(\text{C}/\text{N})$ versus $12+\log(\text{O}/\text{H})$, and $\log(\text{N}/\text{O})$ versus $\log(\text{C}/\text{N})$. We find that both figures confirm the clear linear correlation

found with our observed data for both cases. Interestingly, the main value of $[\text{C}/\text{O}]$ for metal-poor halo stars according to Tomkin et al. (1992) is about -0.60 and about -0.50 according

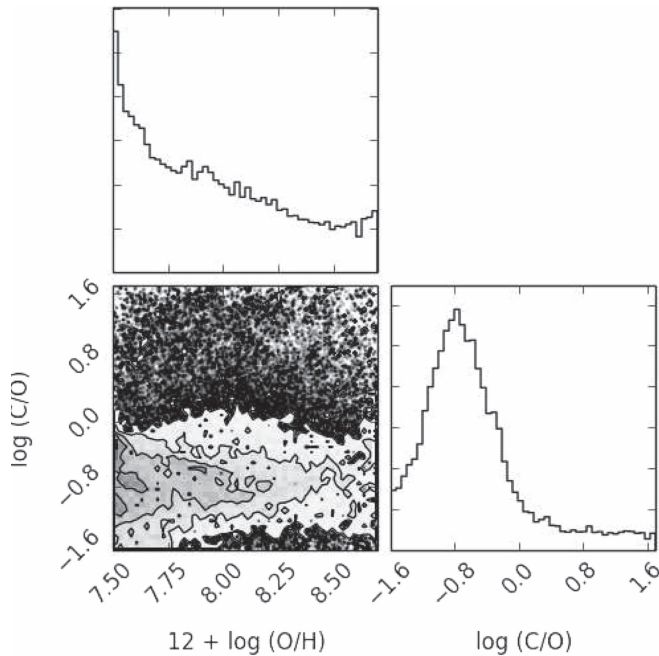


Figure 16. Behavior of the probability of occurrences for the set of abundances $\log(\text{C}/\text{O})$ vs. $12+\log(\text{O}/\text{H})$ for the whole sample. The uppermost panel and the rightmost panel show the histogram of the variable directly below them. This figure includes all of the `CLOUDY` models for all the objects in the sample.

to Fabbian et al. (2009), which translates into a C/O of -0.9 and -0.8 , respectively, adopting the protosolar values for $\text{C}/\text{O}_{\odot} = -0.26 \pm 0.07$ and $12+\log(\text{O}/\text{H})_{\odot} = 8.73 \pm 0.05$ (Asplund et al. 2009). Esteban et al. (2014) determined carbon abundances for a small sample of objects via RLs. They also compared their C abundances to Galactic halo stars. Their C measurements agree with the average C/O from Galactic halo stars shown in Figure 6 of Esteban et al. (2014), also between -0.9 and -0.8 . If our hypothesis of the IMF playing an important role is true, then these plots would again imply that there is a specific behavior for objects that have a similar IMF, and that halo stars are well described by a Kroupa IMF like the one we used for generating the models.


Since intermediate-mass stars contribute to the production of C , the ejection of carbon is delayed with respect to oxygen. Hence, the C/O ratio could be used as an indicator of the relative age of stellar systems (Garnett et al. 1995). Figure 8 in Boyer et al. (2013), which presents the carbon-rich to oxygen-rich (or metal-rich; we will refer to this ratio as C/M) star ratio in the age–metallicity plane for thermally pulsing AGB (TP-AGB) models of Marigo et al. (2013), shows a striking similarity to Figure 16 in the present paper. In this figure, we present the behavior of $\log(\text{C}/\text{O})$ versus $12+\log(\text{O}/\text{H})$ for all of the `CLOUDY` models for the whole sample. If the similarity between these figures is not coincidental, the blue part in the Boyer et al. (2013) figure (corresponding to stars with $\text{C}/\text{O} < 1$) would indicate that the maximum C/M star ratio corresponds to a more positive value of C/O because AGB stars have reached the third dredge-up process and are expelling their carbon into the ISM. As metallicity decreases, it is more difficult to form massive stars, hence there will be more sources of carbon-rich stars.

To test the effect of a different ionizing spectrum on the MCMC `CLOUDY` models, we generated four different synthetic

ionizing spectra and re-ran the MCMC for the least and most metallic objects in our sample, SBS 1415+437 and Tol 9, respectively. The four `Starburst99` ionizing spectra had constant SFR and were taken at an age of 10^8 years. The four models had the following: (i) Geneva stellar evolutionary tracks with rotation velocities of 40% of the break-up velocity and $Z = 0.014$, (ii) the same Geneva stellar evolutionary tracks with $Z = 0.001$, (iii) Geneva stellar evolutionary tracks with rotation velocities of 0% of the break-up velocity and $Z = 0.014$, and (iv) the same Geneva stellar evolutionary tracks with $Z = 0.001$. All other initial `Starburst99` conditions were left as described in Section 4. Spectrum (i) was the same one used to run all models for all objects in the sample. We simply re-ran the chain with the same conditions to see how much the results varied. In order to obtain final abundance values from the MCMC, we used the same average method as previously described in Section 4. We find that there is little variation in the best estimated abundances and the abundances from the models between the 25th and 75th percentiles for all four combination of stellar tracks with metallicity.

As a control experiment to determine the accuracy of the MCMC method we used for abundance verification, we used the line ratios of NGC 5253 presented in Kobulnicky et al. (1997) to run a chain for this object. In this case, we used the abundances given in Kobulnicky et al. (1997) for the H II region as the initial guesses in units of $12+\log(\text{X}/\text{H})$: $\text{He} = 10.90 \pm 0.05$, $\text{O} = 8.19 \pm 0.07$, $\text{C} > 7.46$, $\text{N} = 7.33 \pm 0.13$, $\text{Ne} = 7.32 \pm 0.16$, and $\text{S} = 6.82 \pm 0.17$. The O^{++} and O^+ temperatures they derived are 11,250 and 11,850 K, respectively. Just like with our sample, we initiated the walkers with random positions in the range of the typical observed values. The MCMC diagrams for NGC 5253 show that the models did not converge; however, our sample selection method still worked, and we were able to obtain a first approximation to the abundances. Just as with our sample, we used only the models with $T_{\text{model}} \leq T_{\text{observed}} \pm 2500$ K, and we took an average of these. We determined the uncertainties from the 25th and 75th percentiles, and since the chain did not converge, we expected these uncertainties to be large. This yielded the following abundances in units of $12+\log(\text{X}/\text{H})$: $\text{He}/\text{H} = 10.74 \pm_{0.65}^{0.59}$, $\text{O}/\text{H} = 8.04 \pm 0.23$, $\text{C}/\text{H} = 8.00 \pm_{0.87}^{0.79}$, $\text{N}/\text{H} = 7.03 \pm 0.28$, $\text{Ne}/\text{H} = 7.63 \pm 0.23$, and $\text{S}/\text{H} = 6.32 \pm 0.25$. For those objects where there is no a priori knowledge of the oxygen temperatures, a solution would be to carefully fit a `CLOUDY` model to obtain an approximate of the high- and low-ionization zone temperatures and use them as a proxy for T_{observed} . We then used these center values as the initial condition to re-run the chain. We initiated the walkers with random values but within a tighter “ball” around them (i.e., center value ± 0.2 dex) and ran a chain twice as long (200 runs with 100 walkers). Though the chain did not converge again, the resulting abundances for this re-run were the following: $\text{He}/\text{H} = 10.78 \pm 0.07$, $\text{O}/\text{H} = 8.09 \pm 0.08$, $\text{C}/\text{H} = 8.06 \pm 0.12$, $\text{N}/\text{H} = 7.06 \pm 0.12$, $\text{Ne}/\text{H} = 7.68 \pm 0.11$, and $\text{S}/\text{H} = 6.39 \pm 0.10$, also in units of $12+\log(\text{X}/\text{H})$. These results suggest that our subsample selection method is able to narrow down the set of abundance parameters to obtain a crude first approximation to the abundances even when the chain does not converge, as long as there is a priori knowledge of the high- and low-ionization zone temperatures.

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