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The return to quiescence of Aql X-1 following the 2010 outburst

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ABSTRACT

Aql X-1 is the most prolific low-mass X-ray binary transient hosting a neutron star. In this paper we focus on the return to quiescence following the 2010 outburst of the source. This decay was monitored thanks to 11 pointed observations taken with *XMM-Newton*, *Chandra* and *Swift*. The decay from outburst to quiescence is very fast, with an exponential decay characteristic time-scale of ~ 2 d. Once in quiescence the X-ray flux of Aql X-1 remained constant, with no further signs of variability or decay. The comparison with the only other well-monitored outburst from Aql X-1 (1997) is tale-telling. The luminosities at which the fast decay starts are fully compatible for the two outbursts, hinting at a mechanism intrinsic to the system and possibly related to the neutron star rotation and magnetic field (i.e. the propeller effect). In addition, for both outbursts, the decay profiles are also very similar, likely resulting from the shut-off of the accretion process on to the neutron star surface. Finally, the quiescent neutron star temperatures at the end of the outbursts are well consistent with one another, suggesting a hot neutron star core dominating the thermal balance. Small differences in the quiescent X-ray luminosity among the two outbursts can be attributed to a different level of the power-law component.

Key words: accretion, accretion discs – binaries: close – stars: individual: Aql X-1 – stars: neutron – X-rays: binaries.

1 INTRODUCTION

Bright low-mass X-ray binary transients (LMXTs) hosting a neutron star show two main different outburst phenomenologies and can be broadly divided in ‘quasi-persistent’ and ‘classical’ transients. Quasi-persistent neutron star transients show prolonged active states, lasting up to several years. Sources like MXB 1659–29 (~ 2.5 yr activity, Wijnands et al. 2004; Cackett et al. 2008, and references therein), KS 1731–260 (~ 12.5 yr activity, Wijnands et al. 2002; Cackett et al. 2010, and references therein), EXO 0748–676 (~ 24 yr activity, Degenaar et al. 2009, 2011, and references therein), XTE J1701–462 (~ 1.6 yr, Fridriksson et al. 2010) or HETE J1900.1–2455 (still active, Papitto et al. 2013a, and reference therein) belong to this class. The best sampled outburst X-ray light curve pertains to XTE J1701–462. After a smooth and long

outburst peak, the X-ray luminosity decreased by a factor of ~ 2000 in ~ 13 d, before starting a much slower and smooth decay driven by the crust cooling of the neutron star surface (Fridriksson et al. 2010).

The other broad class of bright transients shows outbursts with a shorter duration (weeks–months), often characterized by a fast rise time (a few days) and a smooth exponential decay (about one month), although many also show more complex outburst light curves. The prototype of this class is Aql X-1 (e.g. Campana, Coti Zelati & D’Avanzo 2013). Monitoring of its 1997 outburst with BeppoSAX showed a steep decay when the luminosity reached a level of $\sim 5 \times 10^{36}$ erg s⁻¹ (at 4.5 kpc), which can be described by an exponential with an e -folding time of ~ 2 d (Campana et al. 1998a). Following this decay, Aql X-1 attained its quiescent level. Similarly steep decays have later been observed (even though with less details) in other sources, such as SAX J1808.4–3658 (Gilfanov et al. 1998; Campana, Stella & Kennea 2008); the Rapid Burster (Masetti et al. 2000) and XTE J1709–267 (Jonker et al. 2003; Degenaar, Wijnands

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& Miller 2013a). This behaviour has been observed also in a few sources classified as faint transients¹ such as IGR J00291+5934 (Jonker et al. 2005), Swift J1756.9–2508 (Krimm et al. 2007) and Swift J1749.4–2807 (Ferrigno et al. 2011; however one outburst of this source showed a linear decay down to quiescence, Wijnands et al. 2009; Campana 2009).

Neutron star transients in quiescence are characterized by an X-ray spectrum made by two different components, one soft and one hard. The soft component can be adequately fitted with a thermal model and is usually interpreted as the result of the cooling of the neutron star surface, heated during repeated outbursts (deep crustal heating; Brown, Bildsten & Rutledge 1998; Wijnands 2011). Alternatively this component could be powered by low-level accretion (van Paradijs et al. 1987; Campana et al. 1998a; Cackett et al. 2010; Fridriksson et al. 2010). The hard component can be described by a power law with photon index $\Gamma \sim 1.5$ –2. Several possibilities have been suggested for its origin, including residual accretion stopped at the magnetospheric radius and shock from a pulsar wind (see e.g. Campana et al. 1998b).

The cooling of the neutron star crust has been studied in detail in quasi-persistent transients, characterized by long cooling times (years), but very little is known on the crust cooling properties of classical (shorter duration) transients. Recently, Degenaar et al. (2013a) monitored the return to quiescence of XTE J1709–267 with *XMM–Newton* observations, finding a very rapid temperature decrease. Two other (faint) transients showed signs of cooling: IGR J17480–2446 in Terzan 5 (Degenaar et al. 2013b) and IGR J17494–3030 (Armas-Padilla, Wijnands & Degenaar 2013).

In this paper we report on the 2010 outburst decay from Aql X-1, the most prolific transient. The 2010 outburst was one of the brightest ever, reaching a peak luminosity of ~ 350 mCrab around 2010 Sept 14. The outburst was preceded by a precursor event. It raised to the maximum in ~ 13 d and then slowly decreased to quiescence in ~ 47 d (Campana et al. 2013). The outburst decay has been followed in great detail with five *XMM–Newton*, four *Chandra* and two *Swift–XRT* observations, monitoring the latest stages of the outburst. In Section 2 we describe the data extraction. In Section 3 we fit the resulting spectra and compare the light-curve decay with that of the only other well-sampled outburst from this source, which occurred in 1997. In Section 4 we discuss the results and in Section 5 we draw our conclusions.

2 DATA REDUCTION

2.1 *XMM–Newton*

Aql X-1 was observed five times with *XMM–Newton* during the decay of the 2010 outburst (see Table 1).² Observations were carried out with the EPIC instrument, which consists of three CCD cameras (one pn and two MOS). The thin filter was used in front of all cameras for all the observations. Data were reprocessed using *SAS* v.11.0. Data were filtered with the `FLAG==0` option and only events with `pattern ≤ 4` (≤ 12) were retained for the pn (MOS) detector. Data were accumulated in the 0.3–10 keV (0.5–10 keV)

¹ Sources in this subclass differ from the bright LMXTs because they reach fainter peak luminosities ($\sim 10^{36}$ erg s⁻¹) and usually show shorter outburst durations.

² One earlier observation was carried out during the same outburst but only the RGS instruments were operated. Due to the limited bandpass we do not consider these data here.

Table 1. Observation log.

| Observation ^a | Start date | Duration (ks) ^b | Obs. ID. | Count rate (c s ⁻¹) |
|--------------------------|--------------|----------------------------|-------------|---------------------------------|
| XMM_1 (BTT) | Oct 12, 2010 | 18.1 | 0085180101 | 14.0 \pm 0.06 |
| XMM_2 (TLT) | Oct 17, 2010 | 6.1 | 0085180201 | 5.48 \pm 0.03 |
| XMM_3 (TLT) | Oct 25, 2010 | 6.4 | 0085180301 | 0.07 \pm 0.01 |
| XMM_4 (FFF) | Oct 29, 2010 | 7.2 | 0085180401 | 0.10 \pm 0.01 |
| XMM_5 (FFF) | Nov 02, 2010 | 14.4 | 0085180501 | 0.10 \pm 0.01 |
| Swift_1 | Oct 15, 2010 | 2.6 | 00030796248 | 0.37 \pm 0.01 |
| Swift_2 | Oct 18, 2010 | 2.8 | 00030796249 | 0.23 \pm 0.01 |
| Chandra_1 | Oct 19, 2010 | 6.4 | 12456 | 0.59 \pm 0.01 |
| Chandra_2 | Oct 22, 2010 | 6.4 | 12457 | 0.26 \pm 0.01 |
| Chandra_3 | Oct 30, 2010 | 6.4 | 12458 | 0.15 \pm 0.01 |
| Chandra_4 | Nov 12, 2010 | 7.0 | 12459 | 0.16 \pm 0.01 |

^aThe observing mode of the three XMM-EPIC instruments is coded as B: burst, T: timing, L: Large window and F: full window. *Swift/XRT* observations were obtained in PC mode. *Chandra* observations were obtained in Timed Exposure mode.

^bExposure times and count rates (pile-up corrected, in the 0.5–10 keV energy range) refer to the MOS1 instrument.

^{*}There are two additional observations with *Swift* that were not used in the present analysis due to the insufficient number of counts for a spectrum.

energy band for the pn (MOS) detector. During all the observations the background remained low and we retained the entire observing time. Burst mode and timing data from the pn detector and timing data from the MOS2 detector were extracted from a 16 pixel wide strip centred on the source (MOS1 timing data were not used because the source fell on a bad column). Large or full window data were extracted from a 750 (650) pixel radius region for the pn (MOS). Count rates for all the observations were reported in Table 1. The background was evaluated from the same CCD in two regions close to Aql X-1 and free of sources and bad columns. The spectral channels were grouped to have at least 20 counts per bin and RMF and ARF files were generated using the latest calibration products.

2.2 *Swift*

Swift observed Aql X-1 several times during the 2010 outburst; however only two observations provided a sufficiently high number of counts to perform an accurate spectral analysis (see Table 1). Both observations were carried out in photon counting (PC) mode with a count rate sufficiently low that no pile-up problems were present. The data were reprocessed using *XRTPIPELINE* v.0.12.6 and filtered with the standard grade scheme, accepting photons with grade in the 0–12 interval. The source counts were extracted from a circular region with radius of 30 and 20 pixels for the two observations, respectively. The background data were extracted from an annulus centred on source with inner and outer radii of 80 and 110 pixels and 70 and 90 pixels, respectively. The ARF files were generated using the task *XRTMKARF* and the appropriate exposure maps. The spectra were binned with at least 20 photons per energy bin in the 0.3–10 keV energy band.

2.3 *Chandra*

Chandra observed Aql X-1 four times during its return to quiescence with the ACIS-S detector in time exposed mode. The exposure time was set for all the observations to 0.44 s per frame (1/8 subarray), in order to mitigate possible pile-up problems (below

Table 2. Spectral fits of all the observations.

| Observation | Column density (10^{21} cm^{-2}) | Power-law photon index | NS temperature (eV) | NS norm. ^a | PL fraction (0.5–10 keV) | χ_{red}^2 (dof) |
|-------------|---|---------------------------|------------------------|-----------------------|-----------------------------|--------------------------------|
| XMM_1 | 5.4 ± 0.1 | $2.06^{+0.02}_{-0.01}$ | <154 | ≤ 1.0 | 100 per cent | 0.93 (1086) |
| Swift_1 | 7.1 ± 1.3 | $2.38^{+0.21}_{-0.19}$ | <270 | ≤ 1.0 | 87 per cent | 0.89 (39) |
| XMM_2 | 3.4 ± 0.1 | $1.16^{+0.16}_{-0.10}$ | 268^{+2}_{-4} | $1.0_{-0.1}$ | 48 per cent | 1.04 (845) |
| Swift_2 | 4.9 ± 1.0 | $0.96^{+1.26}_{-0.99}$ | 249^{+24}_{-32} | $1.0_{-0.3}$ | 41 per cent | 0.91 (25) |
| Chandra_1 | 4.7 ± 0.4 | $1.26^{+0.59}_{-1.03}$ | 212^{+20}_{-10} | $1.0_{-0.4}$ | 48 per cent | 0.95 (120) |
| Chandra_2 | 4.8 ± 0.3 | $0.70^{+0.74}_{-1.03}$ | 180^{+17}_{-4} | $1.0_{-0.3}$ | 36 per cent | 0.81 (55) |
| XMM_3 | 3.5 ± 0.5 | $1.79^{+0.75}_{-0.66}$ | 151^{+6}_{-12} | $1.0_{-0.6}$ | 41 per cent | 0.94 (136) |
| XMM_4 | 4.5 ± 0.3 | $0.93^{+0.64}_{-0.63}$ | 158^{+2}_{-3} | $1.0_{-0.3}$ | 15 per cent | 1.05 (120) |
| Chandra_3 | 4.3 ± 0.9 | $0.49^{+1.61}_{-2.14}$ | 153^{+35}_{-10} | $1.0_{-0.7}$ | 31 per cent | 1.29 (33) |
| XMM_5 | 4.3 ± 0.2 | $1.29^{+0.35}_{-0.35}$ | 153^{+2}_{-2} | $1.0_{-0.2}$ | 21 per cent | 0.99 (212) |
| Chandra_4 | 4.5 ± 0.6 | $0.97^{+1.45}_{-1.61}$ | 157^{+23}_{-23} | $1.0_{-0.5}$ | 30 per cent | 1.39 (35) |

Notes. Errors were computed using $\Delta\chi^2 = 2.71$. We adopted the `TBABS` absorption model and the `NSATMOS` neutron star atmosphere model, while keeping fixed the neutron star mass at $1.4 M_{\odot}$ and radius at 10 km. We adopted a source distance of 4.5 kpc.

Chandra_1 data were affected by pile-up. We fitted the data by including a pile-up model within `XSPEC`. The α parameter was 0.4 ± 0.3 .

^aInitially the neutron star atmosphere model normalization has been fixed to one. In a second step we computed the error on it, leaving the parameter free to vary up to 1.

10 per cent level). In the first observation we find sign of a mild pile-up. For this observation in the spectral fits we included a pile-up model. *Chandra* spectra were extracted using the `CIAO` (4.5) task `SPEXTRACT`. A source region of 8 pixels radius was used and four background regions (radius 15 pixels) surrounding Aql X-1 were selected, avoiding contaminating sources and on the same CCD. Spectra were binned to have at least 20 counts per bin and the analysis was carried out in the 0.5–10 keV energy range. The appropriate response files were generated using the latest calibration release.

3 DATA ANALYSIS

Since the *XMM-Newton* data are those with the highest signal-to-noise ratio, we first used them for selecting the best spectral model. We modelled the Galactic absorption with the `TBABS` model (Wilms, Allen & McCray 2000; using `VERN` cross-sections and `WILM` solar abundances). We tried several single component models never finding an acceptable description of the spectra. A composite model made of a power-law plus a softer component modelled as either a blackbody, a disc model or a neutron star emission model can satisfactorily describe the spectra. However, fits with the disc emission model require an absorption larger than that inferred from the optical reddening (Chevalier et al. 1999). Using a blackbody for the soft component results in an emission radius of only a few km, much smaller than the star radius. This is usually interpreted as the

need for a neutron star atmosphere model which provides more consistent radii (e.g. Zavlin et al. 1996). For these reasons we used in the following the neutron star atmosphere model `NSATMOS` (Heinke et al. 2006) for the soft component.

The absorbed power-law plus `NSATMOS` model provides a good description of the spectra obtained with the different instruments during the Aql X-1 outburst decay (see Table 2). However, the first two observations (XMM_1 and Swift_1) yield a higher column density than the others (and larger than the usual value found in quiescence) and the photon index was close to 2, at variance with that found in all other observations. In addition, in the first *XMM-Newton* observation the temperature of the neutron star atmosphere was very low (actually the thermal component is not required by the data). This might indicate that a spectral model more typical of the outburst phase is necessary for these two observations. Therefore, we fit these spectra with an absorbed cut-off power-law plus atmosphere model. In this way we obtain a harder photon index and more reasonable values for the column density and for the temperature (see the best fit parameters in Table 3).

To characterize in greater depth the variability, we first fit together a set of spectra to see if they can be described by the same column density (in doing this we excluded XMM_1, Swift_1 spectra because they are characterized by a different spectral model, and Chandra_1 spectrum, whose fit includes a pile-up correction model too). The overall fit was good with a reduced $\chi_{\text{red}}^2 = 1.06$ (1524 degrees of freedom, dof) and null hypothesis probability (nhp) of

Table 3. Spectral fits for the first *XMM-Newton* and *Swift* observations.

| Observation | Column density (10^{21} cm^{-2}) | Power-law photon index | Cut-off energy (keV) | NS temperature (eV) | NS norm. | χ_{red}^2 (dof) |
|-------------|---|---------------------------|-------------------------|------------------------|---------------|--------------------------------|
| XMM_1 | 4.8 ± 0.2 | $1.75^{+0.10}_{-0.07}$ | 15^{+7}_{-3} | >259 | $1.0_{-0.60}$ | 0.91 (1084) |
| Swift_1 | 6.3 ± 2.4 | $2.04^{+0.53}_{-1.58}$ | >2 | <270 | ≤ 1.0 | 0.91 (38) |

Notes. Errors were computed using $\Delta\chi^2 = 2.71$. We adopted the `TBABS` absorption model and the `NSATMOS` neutron star atmosphere model, keeping fixed the neutron star mass at $1.4 M_{\odot}$ and the radius at 10 km. We adopted a source distance of 4.5 kpc.

Table 4. Spectral fits of all the observations with the same column density for all the observations.

| Observation | Column density (10^{21} cm^{-2}) | Power-law photon index | NS temperature (eV) | Unabs. luminosity (0.5–10 keV cgs) | χ_{red}^2 (dof) |
|-------------|---|---------------------------|------------------------|---------------------------------------|--------------------------------|
| XMM_1 | Fixed | $1.46^{+0.03}_{-0.03}$ | >270 | 1.0×10^{36} | 0.91 (1085) |
| Swift_1 | Fixed | $1.19^{+0.34}_{-0.66}$ | 8 – 270 | 1.9×10^{35} | 1.05 (38) |
| XMM_2 | 4.0 ± 0.1 | $1.83^{+0.07}_{-0.09}$ | 258^{+5}_{-5} | 7.4×10^{34} | 1.06 (1524) |
| Swift_2 | – | $0.99^{+0.73}_{-0.45}$ | 253^{+11}_{-22} | 4.2×10^{34} | |
| Chandra_1 | Fixed | $1.05^{+0.33}_{-0.42}$ | 208^{+3}_{-4} | 3.8×10^{33} | 1.16 (122) |
| Chandra_2 | – | $1.02^{+0.57}_{-0.54}$ | 177^{+3}_{-5} | 7.1×10^{33} | |
| XMM_3 | – | $2.64^{+0.15}_{-0.29}$ | 149^{+8}_{-7} | 5.3×10^{33} | |
| XMM_4 | – | $0.51^{+0.80}_{-0.72}$ | 160^{+1}_{-2} | 3.9×10^{33} | |
| Chandra_3 | – | $0.36^{+1.39}_{-1.08}$ | 160^{+3}_{-6} | 4.7×10^{33} | |
| XMM_5 | – | $1.18^{+0.38}_{-0.33}$ | 159^{+1}_{-2} | 3.8×10^{33} | |
| Chandra_4 | – | $1.14^{+1.22}_{-1.31}$ | 160^{+4}_{-14} | 4.6×10^{33} | |

XMM_1 was fitted with a cut-off power law with $E_c = 7.5^{+0.7}_{-0.6}$ keV.

Table 5. Spectral fits of all the observations with the same column density and power-law index for all the observations.

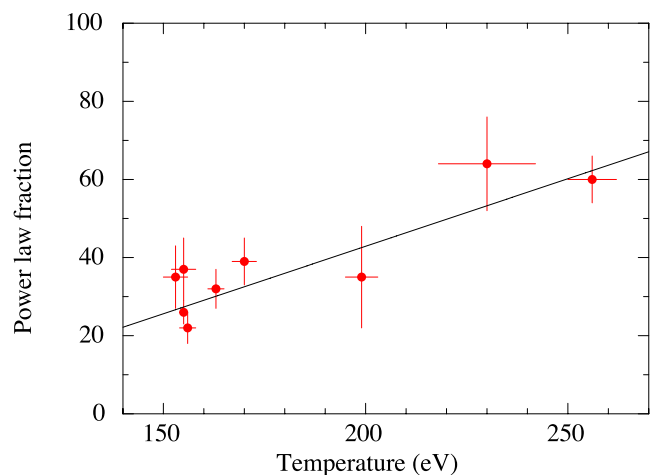
| Observation | Column density (10^{21} cm^{-2}) | Power-law photon index | NS temperature (eV) | Power-law fraction | χ_{red}^2 (dof) |
|-------------|---|---------------------------|------------------------|-----------------------|--------------------------------|
| XMM_2 | 3.8 ± 0.1 | $1.70^{+0.07}_{-0.13}$ | 265^{+6}_{-5} | 60 ± 6 per cent | 1.08 (1531) |
| Swift_2 | – | – | 230^{+12}_{-12} | 62 ± 12 per cent | |
| Chandra_1 | Fixed | Fixed | 199^{+3}_{-4} | 35 ± 13 per cent | 1.46 (123) |
| Chandra_2 | – | – | 170^{+3}_{-3} | 39 ± 6 per cent | |
| XMM_3 | – | – | 163^{+2}_{-2} | 32 ± 5 per cent | |
| XMM_4 | – | – | 156^{+2}_{-2} | 22 ± 4 per cent | |
| Chandra_3 | – | – | 153^{+3}_{-3} | 35 ± 8 per cent | |
| XMM_5 | – | – | 155^{+1}_{-1} | 26 ± 3 per cent | |
| Chandra_4 | – | – | 155^{+3}_{-3} | 37 ± 8 per cent | |

XMM_1 and Swift_1 were fitted with a cut-off power law.

0.04. Results of the fit are reported in Table 4. In Table 3 we also include the fits for XMM_1, Swift_1 spectra, obtained by fixing the column density to the best common fit found before. We then tried to group another spectral parameter. If we fit all the spectra in Table 4 by keeping the same neutron star temperature, we end up with a $\chi_{\text{red}}^2 = 1.29$ (1531 dof), with an nhp of $\sim 10^{-13}$ (see Table 5). If, instead, we keep the same power-law photon index we obtain a good fit with $\chi_{\text{red}}^2 = 1.08$ (1531 dof, nhp=0.02). As a further investigation, we computed the power-law fraction to the total flux and found that it is positively correlated to the neutron star temperature, at least during the decay phase (see Fig. 1).

3.1 Light-curve evolution

Based on the spectral fits reported in Section 3 we derived the evolution of the outburst decay shown in Fig. 2, where we plot the 0.5–10 keV (unabsorbed) luminosity (we adopt a distance of 4.5 kpc; Galloway et al. 2008). To illustrate the initial part of the outburst we used data from the MAXI satellite³ (Matsuoka et al. 2009), converted into luminosity assuming a Crab-like spectrum and a column density of $4 \times 10^{21} \text{ cm}^{-2}$. The overall light curve can be divided into three parts: a shallow decay followed by a steep decrease spanning three orders of magnitude in ~ 10 d and finally a steady level. We modelled these three time intervals with

**Figure 1.** Correlation between the power-law fraction and the neutron star atmosphere temperature during the 2010 outburst based on spectral fits, taken from Table 5.

a Gaussian, an exponential and a constant, respectively starting from the outburst peak. The Gaussian width (σ) was 15 ± 1 d and the steep decay started before the first XMM–Newton observation (around MJD 55479) and had an e -folding time of 1.7 ± 0.1 d. The luminosity at the knee was $5.9^{+0.2}_{-0.8} \times 10^{36} \text{ erg s}^{-1}$ (the error has been computed by propagating the error on the knee time into

³ The light curve has been obtained from <http://maxi.riken.jp>

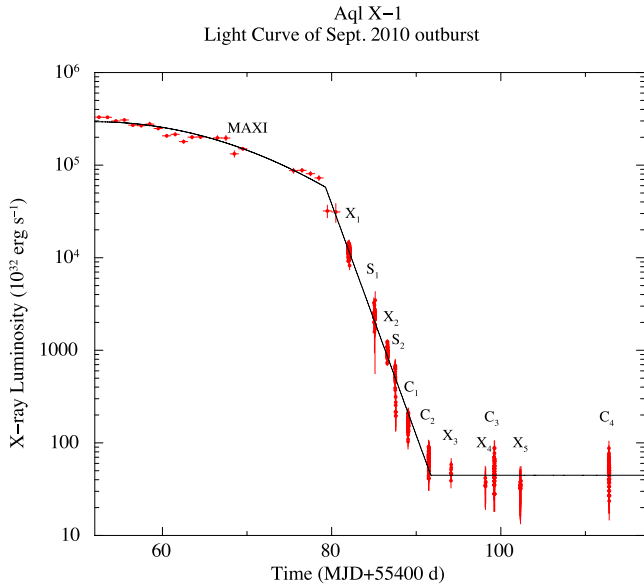


Figure 2. X-ray luminosity light curve of the 2010 outburst from Aql X-1 as observed by *MAXI*, *XMM-Newton*, *Swift* and *Chandra*. Count rates were converted into luminosities based on the results of the spectral fits and a distance of 4.5 kpc. The *MAXI* light curve is binned to 1 d, and the focusing instruments light curves are binned to 50 s.

the luminosity curve). Quiescence started at MJD 55491, reaching a level of $4.5 \times 10^{33} \text{ erg s}^{-1}$. This was computed by assuming the same spectrum for the last four observations (two *XMM-Newton* and two *Chandra*). The fit was fairly good with $\chi_{\text{red}}^2 = 1.1$ (418 dof). The best-fitting values were $N_{\text{H}} = (4.4 \pm 0.1) \times 10^{21} \text{ cm}^{-2}$, $T = 154 \pm 2 \text{ eV}$ and $\Gamma = 1.3 \pm 0.3$. The power law comprised ~ 20 per cent of the total quiescent flux.

To compare the outburst light curve with that observed in 1997 with *BeppoSAX* and *RossixTE* (Campana et al. 1998b), we reanalysed the data from these satellites using the same spectral models and distance adopted for the 2010 observations. The 1997 outburst is the only one with an adequately sampled decay light curve. As shown in Fig. 3, the temporal evolution of the 1997 outburst was very similar to that of the 2010 outburst. It can be described by a Gaussian with $\sigma = 15 \pm 1 \text{ d}$ followed by an exponential decay with e -folding time of $1.8 \pm 0.1 \text{ d}$ (starting at MJD 50513). The luminosity at the knee was $5.2^{+0.2}_{-0.9} \times 10^{36} \text{ erg s}^{-1}$. Considering the uncertainties involved in the count rate to flux conversion and on the light curve modelling, this value is consistent with that observed in 2010.

The quiescence after the 1997 outburst started at MJD 50521, at a level of $5.9 \times 10^{33} \text{ erg s}^{-1}$,⁴ and the spectral analysis gave a neutron star temperature of $159^{+10}_{-6} \text{ eV}$, similar to that observed in 2010. The higher quiescent luminosity (a factor of 1.3 compared to that of 2010) can be accounted by the power-law component ($\Gamma = 1.4 \pm 0.6$), which comprised ~ 40 per cent of the total quiescent flux, being the temperatures of the two quiescent states comparable. Campana et al. (2013) estimated the energetics of the two outbursts by integrating the overall light curves (as observed by *RXTE/ASM* in 1997) and *MAXI* (for the 2010 outburst), obtaining $\sim 2.2 \times 10^{43} \text{ erg}$

⁴ This luminosity is slightly different from that reported in Campana et al. (1998b) because we adopted a neutron star atmosphere model instead of a blackbody, a column density fixed to $4.4 \times 10^{21} \text{ cm}^{-2}$, and a distance of 4.5 kpc.

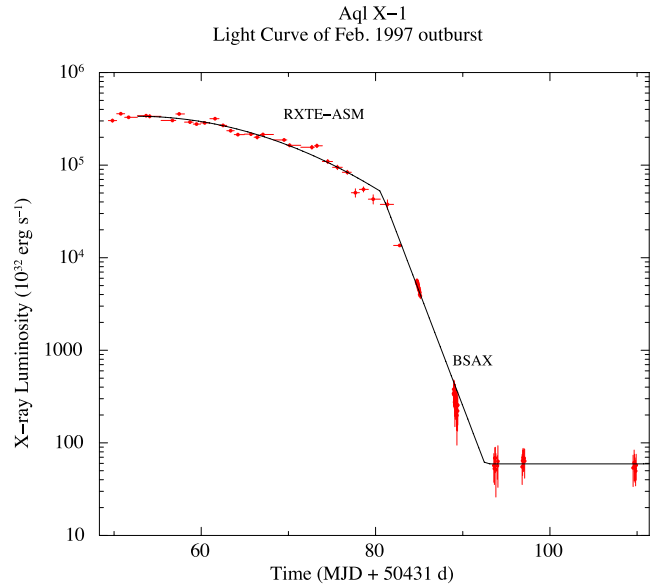


Figure 3. X-ray luminosity light-curve evolution of the 1997 outburst from Aql X-1 as observed by *RossixTE* and *BeppoSAX*. Count rates were converted into luminosities based on the results of the spectral fits and a distance of 4.5 kpc.

Table 6. Comparison among the 2010 and 1997 Aql X-1 outbursts.

| | 2010 | 1997 |
|---|------------------------------------|------------------------------------|
| Total energetics (erg) | $\sim 2.8 \times 10^{43}$ | $\sim 2.2 \times 10^{43}$ |
| Duration (d) | ~ 60 | ~ 90 |
| Knee Lum. (erg s^{-1}) | $5.9^{+0.2}_{-0.8} \times 10^{36}$ | $5.2^{+0.2}_{-0.9} \times 10^{36}$ |
| Decay e -fold time (d) | 1.7 ± 0.1 | 1.8 ± 0.1 |
| Quiescent Lum. ($10^{33} \text{ erg s}^{-1}$) | (4.5 ± 0.2) | (5.9 ± 0.9) |
| Quiescent temperature (eV) | 158^{+1}_{-1} | 159^{+10}_{-6} |

and $\sim 2.8 \times 10^{43} \text{ erg}$, respectively. The ratio of the outbursts' fluences (2010/1997) is a factor of ~ 1.3 . The main parameters derived for the two outbursts are summarized in Table 6.

3.2 Temperature evolution

The temperatures derived with the spectral analysis described above allow us to derive a thermal evolution curve for the outburst decay. The derived temperatures were first converted into temperatures observed by a distant observer as $kT^\infty = kT/(1+z)$, where $1+z = (1 - R_s/R)^{-1/2} = 1.31$ is the gravitational redshift factor for $M = 1.4M_\odot$ and $R = 10 \text{ km}$ (with $R_s = 2GM/c^2$ being the Schwarzschild radius, G the gravitational constant and c the speed of light). Cooling curves are usually measured from the end of the outburst, which in our case cannot be assessed easily. We decided to start from the knee marking the onset of the steep decay (MJD 55479). The resulting temporal evolution of the neutron star temperature is shown in Fig. 4. We tried to fit it with either a power law ($at^{-\alpha} + b$) or an exponential decay [$a \exp(-t/\tau) + b$]. In both cases we included baseline constant level (b) in the fit. The power-law fit gives a reduced $\chi_{\text{red}}^2 = 3.3$ with 6 dof with $\alpha = 3.5 \pm 0.4$ and $b = 119 \pm 2 \text{ eV}$. A better description of the data is obtained with the exponential decay, which gives $\tau = 2.5 \pm 0.3 \text{ d}$ and $b = 121 \pm 1 \text{ eV}$ ($\chi_{\text{red}}^2 = 2.6$ for 6 dof). These values are markedly different from what is usually found for quasi-persistent LMXBs, where the

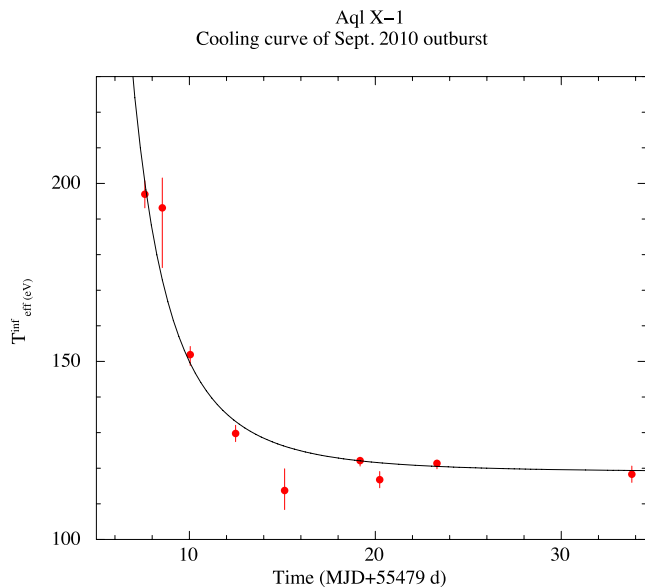


Figure 4. Temperature evolution of the neutron star component during the 2010 outburst, based on spectral modelling with column density and power-law spectral slope tied across all the observations. Temperature values are taken from Table 5 and corrected for the gravitational redshift.

fit with an exponential decay returns an e -folding time-scale in the range 200–800 d (Degenaar et al. 2013b, and references therein). On an observational basis, it is apparent that the decay of Aql X-1 is much faster/steeper.

4 DISCUSSION

4.1 Luminosity drop off

The steep flux decay observed in the outburst of X-ray transients is still of unknown origin. It is important to underline that such a steep decay occurs in a number of different systems. It has been observed not only in neutron star transients but also in X-ray binaries hosting a black hole (e.g. Homan et al. 2013, see also Shahbaz, Charles & King 1998), as well as in cataclysmic variables (WZ Sge type; Kato et al. 2004). Given its ubiquity, it could be related to the physics of the accretion disc rather than to the interaction of the magnetic field with the accreting matter (e.g. centrifugal inhibition of accretion, a.k.a. propeller effect). The disc instability model (DIM; Lasota 2001) is the leading model to explain LMXT (as well as dwarf novae) outbursts: an accretion disc is stable if hydrogen is everywhere ionized. If instead its temperature is low enough for hydrogen to recombine, the disc becomes thermally and viscously unstable. In this regime the disc oscillates between a hot, ionized state (outburst) and a cold, neutral state (quiescence), due to the strong dependence of the opacity on temperature when hydrogen is partially ionized (Lasota 2001).

DIM predicts exponential-like decays as long as the outer disc temperature is higher than a critical temperature; it is expected to revert to a linear decay thereafter. This transition in the X-ray light curve is also called ‘brink’ (King & Ritter 1998). The presence of a luminosity brink was reported in several Aql X-1 outbursts (Campana et al. 2013). Campana et al. (2013) found a variable outer disc radius (a factor of ~ 5), resulting in different brink luminosities on an outburst-by-outburst basis. In addition to the X-ray luminosity brink, Asai et al. (2013) claimed a further transition in the decay

light curve of Aql X-1 and 4U 1608–52, which they attribute to the propeller effect (no predictions for the outburst decay in case of a propeller effect have ever been made). In the case of Aql X-1 they estimated a luminosity of $\sim 8 \times 10^{35} \text{ erg s}^{-1}$ at propeller onset.

The first part of the 2010 outburst of Aql X-1, from the peak onward, can be successfully modelled with a Gaussian. A knee is present at a luminosity of $5.9^{+0.2}_{-0.7} \times 10^{36} \text{ erg s}^{-1}$, followed by an exponential decay with an e -folding time of $1.7 \pm 0.1 \text{ d}$. While the light curve before the knee could also be described with an exponential, a linear fit after the knee, as predicted in the DIM if we interpret the knee as the brink luminosity, is clearly rejected by the data. In addition, the data do not reveal any departure from an exponential decay, at variance with Asai et al. (2013). On these basis we consider unlikely that the steep decay is due to an outer disc transition. This also indicates that outburst light curves probed by all-sky monitoring instruments are difficult to model when the transition is close to the instrumental detection limit.

The 1997 outburst of Aql X-1 was monitored with sufficient detail to estimate the knee luminosity (Campana et al. 1998b). By considering the uncertainties involved by the different X-ray instruments, the 1997 and 2010 knee luminosities were very similar and statistically compatible, with a mean 0.5–10 keV luminosity of $(5.5 \pm 0.3) \times 10^{36} \text{ erg s}^{-1}$. Given the variations observed, e.g., in the brink luminosity of other outbursts of Aql X-1 (Campana et al. 2013), it is puzzling that the same knee luminosity is observed in the two best-studied outbursts (in addition a number of outbursts show a brink luminosity consistent with this value). One possibility is that the propeller effect due to the neutron star magnetic field plays a role, through the action of the centrifugal barrier. Assuming a relation for the onset of the propeller linking the knee luminosity to the neutron star spin period and magnetic field as in Asai et al. (2013), one can estimate a magnetic field of $B \sim (1 - 4) \times 10^8 \text{ G}$ (including a factor of 2 uncertainty in the Alfvén radius definition). Thus, we are observing the shut-off of the accretion process on to the neutron star surface. One possibility is that a ‘competition’ occurs among the propeller onset and the (variable) brink onset, depending on the outburst evolution (e.g. irradiation properties, mass inflow rate, etc.). In the case of black holes, or (low-magnetic field) white dwarfs, only the disc drives the luminosity drop off. This conjecture can be tested by doing similar in depth studies of the decay behaviour in black hole transients.

4.2 Temperature evolution

The temperature evolution observed during the light-curve decay was quite fast. The temperature variations could be fitted with an exponential decay with $\tau = 2.5 \pm 0.3 \text{ d}$. This is somewhat slower than the luminosity decay ($1.7 \pm 0.1 \text{ d}$), but still extremely fast. In quasi-persistent transients the decay is much longer in comparison with $\tau \gtrsim 100 \text{ d}$. Recently, the return to quiescence of a few classical neutron star transients has been monitored in detail. IGR J17480–2446 in Terzan 5 exhibited an 11-week accretion episode (to be compared with ~ 9 weeks of the 2010 Aql X-1 outburst) and a slow neutron star cooling ($\sim 800 \text{ d}$; Degenaar et al. 2013b). XTE J1709–267 showed a very rapid neutron star temperature decrease following an approximately two months outburst. In particular, the temperature decreased from $\sim 160 \text{ eV}$ to $\sim 150 \text{ eV}$ in 8 hr and an exponential decay with $\tau \sim 5.6 \pm 1.9 \text{ d}$ (Degenaar et al. 2013a). IGR J17494–3030 showed a very short 10-d outburst. During an *XMM-Newton* observation the neutron star temperature decreased, with an e -folding time-scale of $\sim 3 \text{ d}$ (Armas-Padilla et al. 2013). Even though the fast decay of XTE J1709–267 has been

interpreted as emission from a cooling layer in the neutron star atmosphere (Degenaar et al. 2013a), the very fast cooling time of IGR J17494–3030 likely hints to accretion as the powering source of the observed X-ray luminosity. We note that the X-ray spectrum arising from a low-level accreting X-ray source is very similar to that of a cooling neutron star atmosphere (Zampieri et al. 1995; Campana, Mereghetti & Sidoli 1997). Given that the X-ray luminosity resulting from accretion on to the neutron star surface is expected to track the mass accretion rate, in this picture the luminosity drop off should be caused by a drastic reduction of mass accretion rate. As discussed above, this may result from a low-viscosity disc setting in or from the propeller onset. In the second case we do not expect a sudden and complete halting of the mass inflow rate, but rather a fast drop of the X-ray luminosity. One possibility is therefore that we are observing the closure of the centrifugal barrier. In this case, the cooling of the neutron star surface is overwhelmed by low-level accretion.

A rapid drop off of the X-ray luminosity has also been observed in the quasi-persistent source XTE J1701–462 (Fridriksson et al. 2010). This source remained active for ~ 600 d and turned off in about 13 d. Its phase exponential decay, with $\tau \sim 1.2$ d (Fridriksson et al. 2010), has been interpreted as due to the cessation of the accretion process. Being heated during a very long outburst, it is not surprising that the neutron star surface emission set in at a high X-ray luminosity $\sim 10^{34}$ erg s $^{-1}$ and showed a slow cooling with exponential decay of ~ 120 d. Quasi-persistent transient sources as well as XTE J1701–462 experienced long outburst durations, which likely brought their crusts out of thermal equilibrium with the core. The crust reached the equilibrium with the core on a long time-scale (years, see Brown et al. 1998, for details). Due to the shorter outburst durations and the higher quiescent luminosity, this is not observed in Aql X-1.

4.3 Quiescent level

As discussed above, the quiescent state of Aql X-1 after the 2010 outburst did not show any sign of further decay (at variance with quasi-persistent low-mass transients). In addition, the quiescent luminosity is slightly lower than that observed immediately following the 1997 outburst (factor ~ 1.3). The 1997 and 2010 outburst fluences differ instead by a factor of ~ 1.3 in the opposite sense. This can be accounted for by the hard luminosity differences and might indicate that in quiescence we are observing the true core temperature which is in thermal equilibrium with the atmosphere. In fact, even if the 1997 and 2010 outbursts were among the strongest the Aql X-1 records in terms of fluence, the core temperature is thought to be largely unaffected on short (years) time-scales, i.e. between different outbursts (Brown et al. 1998; Colpi et al. 2001). Being the heat deposited during outbursts small and the core already hot only a small cooling might be expected. Unfortunately, the quiescence of Aql X-1 is restless, showing flares (ten-fold increase, Coti Zelati et al. 2013; see also Bernardini et al. 2013), and this slow cooling may be difficult to probe. In addition, the frequent recurrence of the outbursts hampers the possibility to observe the post-outburst behaviour after individual events for long time interval. Consistently, Aql X-1 is one of the few sources lying on the prediction of standard cooling (Heinke et al. 2010; Wijnands, Degenaar & Page 2013). Differences in quiescent levels following different outbursts could perhaps indicate different levels of quiescent accretion (if at work) or different interaction of a turned on millisecond pulsar with the surrounding medium (Campana et al. 1998a; Papitto et al. 2013b), both likely involving the power-law hard spectral compo-

nent, or a different composition of the neutron star atmosphere at the end of the outburst resulting in different observed luminosity for the same core temperature (Brown et al. 1998; Degenaar et al. 2013b).

5 CONCLUSIONS

We monitored the decay to quiescence of the Aql X-1 2010 outburst. Our data set, based on five *XMM-Newton*, four *Chandra* and two *Swift* observations, provides the most detailed view of a neutron star binary transient outburst decay. The outburst decay shows a pronounced steepening when the X-ray luminosity drops below 5×10^{36} erg s $^{-1}$ (at a 4.5 kpc distance). Afterwards, the luminosity drops exponentially with an e -folding time of ~ 2 d, reducing its luminosity by a factor of ~ 1000 in ~ 10 d. We were also able to track the evolution of the temperature of the soft spectral component characterizing the X-ray spectrum. The temperature decrease is quite fast too, with an e -folding time of ~ 3 d.

Comparing the Aql X-1 2010 outburst decay to the 1997 outburst decay (Campana et al. 1998a), we find the same knee luminosity and the same fast decay. This likely calls for the same mechanism responsible for the outburst turn off. We suggest that this occurs due to the propeller inhibition of accretion, induced by the fast spinning neutron star magnetosphere. The steep luminosity decay would then be due to the closure of the centrifugal barrier, with the temperature decay tracking the switch off of the mass inflow on to the neutron star surface.

The quiescent luminosities at the end of the 1997 and 2010 outbursts are somewhat different; the quiescent neutron star temperatures are instead well consistent with one another. This suggests that a hot neutron star core is dominating the thermal balance and X-ray luminosity differences can be attributed to a different level of the quiescent power-law component.

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REFERENCES

- Armas-Padilla M., Wijnands R., Degenaar N., 2013, *MNRAS*, 436, L89
 Asai K. et al., 2013, *ApJ*, 773, 117
 Bernardini F., Cackett E. M., Brown E. F., D'Angelo C., Degenaar N., Miller J. M., Reynolds M., Wijnands R., 2013, *MNRAS*, 436, 2465
 Brown E. F., Bildsten L., Rutledge R. E., 1998, *ApJ*, 504, L95
 Cackett E. M., Wijnands R., Miller J. M., Brown E. F., Degenaar N., 2008, *ApJ*, 687, L87
 Cackett E. M., Brown E. F., Cumming A., Degenaar N., Miller J. M., Wijnands R., 2010, *ApJ*, 722, L137
 Campana S., 2009, *ApJ*, 699, 1144
 Campana S., Mereghetti S., Sidoli L., 1997, *A&A*, 32, 783
 Campana S., Stella L., Mereghetti S., Colpi M., Tavani M., Ricci D., Dal Fiume D., Belloni T., 1998a, *ApJ*, 499, L65
 Campana S., Colpi M., Mereghetti S., Stella L., Tavani M., 1998b, *A&AR*, 8, 279
 Campana S., Stella L., Kennea J. A., 2008, *ApJ*, 684, L99
 Campana S., Coti Zelati F., D'Avanzo P., 2013, *MNRAS*, 432, 1695
 Chevalier C., Ilovaisky S. A., Leisy P., Patat F., 1999, *A&A*, 347, L51
 Colpi M., Geppert U., Page D., Possenti A., 2001, *ApJ*, 548, L175
 Coti Zelati F., Campana S., D'Avanzo P., Melandri A., 2013, *MNRAS*, 438, 2634

- Degenaar N. et al., 2009, *MNRAS*, 396, L26
 Degenaar N. et al., 2011, *MNRAS*, 412, 1409
 Degenaar N., Wijnands R., Miller J. M., 2013a, *ApJ*, 767, L31
 Degenaar N. et al., 2013b, *ApJ*, 775, 48
 Ferrigno C. et al., 2011, *A&A*, 525, A48
 Fridriksson J. K. et al., 2010, *ApJ*, 714, 270
 Galloway D. K., Muno M. P., Hartman J. M., Psaltis D., Chakrabarty D., 2008, *ApJS*, 179, 360
 Gilfanov M., Revnivtsev M., Sunyaev R., Churazov E., 1998, *A&A*, 338, L83
 Heinke C. O., Rybicki G. B., Narayan R., Grindlay J. E., 2006, *ApJ*, 644, 1090
 Heinke C. O. et al., 2010, *ApJ*, 714, 894
 Homan J., Fridriksson J. K., Jonker P. G., Russell D. M., Gallo E., Kuulkers E., Rea N., Altamirano D., 2013, *ApJ*, 775, 9
 Jonker P. G., Méndez M., Nelemans G., Wijnands R., van der Klis M., 2003, *MNRAS*, 341, 823
 Jonker P. G., Campana S., Steeghs D., Torres M. A. P., Galloway D. K., Markwardt C. B., Chakrabarty D., Swank J., 2005, *MNRAS*, 361, 511
 Kato T., Nogami D., Matsumoto K., Baba H., 2004, *PASJ*, 56, 109
 King A. R., Ritter H., 1998, *MNRAS*, 293, L42
 Krimm H. A. et al., 2007, *ApJ*, 668, L147
 Lasota J.-P., 2001, *New Astron. Rev.*, 45, 449
 Masetti N. et al., 2000, *A&A*, 363, 188
 Matsuoka M. et al., 2009, *PASJ*, 61, 999
 Papitto A. et al., 2013a, *MNRAS*, 429, 3411
 Papitto A. et al., 2013b, *Nature*, 501, 517
 Shahbaz T., Charles P. A., King A. R., 1998, *MNRAS*, 301, 382
 van Paradijs J., Verbunt F., Shafer R. A., Arnaud K. A., 1987, *A&A*, 182, 47
 Wijnands R., 2011, *proc. Fast X-ray Timing and Spectroscopy at Extreme Count Rates (HTRS 2011)*. Champéry, Switzerland
 Wijnands R., Guainazzi M., van der Klis M., Méndez M., 2002, *ApJ*, 573, L45
 Wijnands R., Homan J., Miller J. M., Lewin W. H. G., 2004, *ApJ*, 606, L61
 Wijnands R., Rol E., Cackett E., Starling R. L. C., Remillard R. A., 2009, *MNRAS*, 393, 126
 Wijnands R., Degenaar N., Page D., 2013, *MNRAS*, 432, 2366
 Wilms J., Allen A., McCray R., 2000, *ApJ*, 542, 914
 Zampieri L., Turolla R., Zane S., Treves A., 1995, *ApJ*, 439, 849
 Zavlin V. E., Pavlov G. G., Shibanov Yu. A., 1996, *A&A*, 315, 141

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