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**DOI**

[10.1051/0004-6361/201014037e](https://doi.org/10.1051/0004-6361/201014037e)

**Publication date**

2011

**Document Version**

Final published version

**Published in**

Astronomy & Astrophysics

[Link to publication](#)

**Citation for published version (APA):**

Baudin, F., Barban, C., Belkacem, K., Hekker, S., Morel, T., Samadi, R., Benomar, O., Goupil, M. J., Carrier, F., Ballot, J., Deheuvels, S., de Ridder, J., Hatzes, A. P., Kallinger, T., & Weiss, W. W. (2011). Amplitudes and lifetimes of solar-like oscillations observed by CoRoT: red-giant versus main-sequence stars. (Corrigendum). *Astronomy & Astrophysics*, 535, 1. <https://doi.org/10.1051/0004-6361/201014037e>

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# Amplitudes and lifetimes of solar-like oscillations observed by CoRoT

## Red-giant versus main-sequence stars

### (Corrigendum)

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A&A, 529, A84 (2011), DOI: 10.1051/0004-6361/201014037

**Key words.** stars: oscillations – methods: data analysis – asteroseismology – errata, addenda

In Sects. 4.3 and 4.4, uncertainties were not properly taken into account in the computation of the scaling relation between mode widths (Sect. 4.3) and effective temperature for main-sequence stars (MS), and of the scaling relation between mode amplitude and effective temperature, again for main-sequence stars (other scaling relations are correct). The correct results are

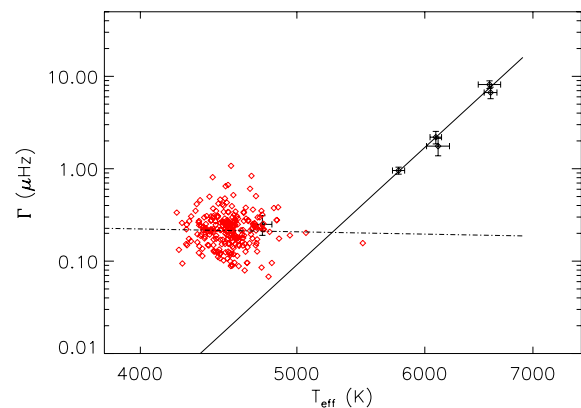
$$\Gamma \propto T_{\text{eff}}^s \quad \text{with} \quad \begin{cases} s \approx -0.3 \pm 0.9 & \text{for RG (unchanged)} \\ s \approx 16 \pm 2 & \text{for MS (corrected)} \end{cases}$$

and:

$$A_{\text{max}}^I \propto \left(\frac{L}{M}\right)^s \quad \text{with} \quad \begin{cases} s = 0.92 \pm 0.14 & \text{for RG (unchanged)} \\ s = 0.42 \pm 0.14 & \text{for MS (corrected)}. \end{cases}$$

The corrected values are only slightly changed and well inside former uncertainties. Moreover, the corrected uncertainties are smaller than the former ones. Figure 7 is also slightly changed (see Fig. 1). The conclusions remain unchanged with these more precise results.

*Acknowledgements.* This wrong computation was pointed out by T. Appourchaux who performed a thorough check of the former published results.



**Fig. 1.** Measured mode linewidths versus  $T_{\text{eff}}$  for the red giants ( $T_{\text{eff}} < 5000$  K) and for the main-sequence stars ( $T_{\text{eff}} > 5000$  K). The dot-dash line indicates the fit to the red-giant width and the triple-dot-dash line the fit to main-sequence stars.