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SCUBA2 HIGH REDSHIFT BRIGHT QUASAR SURVEY: FAR-INFRARED PROPERTIES AND WEAK-LINE FEATURES

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ABSTRACT

We present a submillimeter continuum survey (‘SCUBA2 High rEdshift bRight quasaR surveY’, hereafter SHERRY) of 54 high-redshift quasars at $5.6 < z < 6.9$ with quasar bolometric luminosities in the range of $0.2\text{--}5 \times 10^{14} L_{\odot}$, using the Submillimetre Common-User Bolometer Array-2 (SCUBA2) on the James Clerk Maxwell Telescope. About 30% (16/54) of the sources are detected with a typical $850\mu\text{m}$ rms sensitivity of $1.2 \text{ mJy beam}^{-1}$ ($S_{\nu, 850\mu\text{m}} = 4\text{--}5 \text{ mJy}$, at $> 3.5\sigma$). The new SHERRY detections indicate far-infrared (FIR) luminosities of 3.5×10^{12} to $1.4 \times 10^{13} L_{\odot}$, implying extreme star formation rates of $90\text{--}1060 M_{\odot} \text{ yr}^{-1}$ in the quasar host galaxies. Compared with $z = 2\text{--}5$ samples, the FIR-luminous quasars ($L_{\text{FIR}} > 10^{13} L_{\odot}$) are more rare at $z \sim 6$. The optical/near-infrared (NIR) spectra of these objects show that 11% (6/54) of the sources have weak Ly α emission-line features, which may relate to different sub-phases of the central active galactic nuclei (AGNs). Our SCUBA2 survey confirms the trend reported in the literature that quasars with submillimeter detections tend to have weaker ultraviolet (UV) emission lines compared to quasars with non-detections. The connection between weak UV quasar line emission and bright dust continuum emission powered by massive star formation may suggest an early phase of AGN-galaxy evolution, in which the broad-line region is starting to develop slowly or shielded from the central ionization source, and has unusual properties such as weak-line features or bright FIR emission.

Keywords: cosmology: observations – quasars: general – galaxies: active – galaxies: high redshift – submillimeter: galaxies

1. INTRODUCTION

Quasars probe the rapid accreting phase of the supermassive black holes (SMBHs) in the center of galaxies. Quasar surveys toward the highest redshift open

a unique window for the studies of SMBH and galaxy evolution in the early universe. The first quasar at $z \sim 6$ was discovered by Fan et al. (2000) and currently there are more than 250 $z > 5.6$ quasars known from the optical and near-infrared (NIR) surveys, such as the Sloan Digital Sky Survey (SDSS; e.g., Fan et al. 2006; Jiang et al. 2008, 2015, 2016), CanadaFrance High- z Quasar Survey (CFHQS; e.g., Willott et al. 2007, 2010, UKIRT Infrared Deep Sky Survey (UKIDSS; e.g., Venemans

et al. 2007; Mortlock et al. 2011; Wang et al. 2019), VISTA Kilo-degree Infrared Galaxy (VIKING; e.g., Venemans et al. 2013, 2015a), VLT Survey Telescope-ATLAS (VST-ATLAS; Carnall et al. 2015), Dark Energy Survey (DES; Reed et al. 2015), Hyper Suprime-Cam (HSC; Matsuoka et al. 2016; Matsuoka et al. 2018a,b, 2019), and Panoramic Survey Telescope and Rapid Response System (PS1; Bañados et al. 2014; Venemans et al. 2015b; Bañados et al. 2016). This large sample allows us to study the early growth of the SMBHs and galaxies close to cosmic reionization in the systems with a wide range of SMBH masses and quasar luminosities; as well as probe the redshift evolution of SMBHs and host galaxies.

Observations at submillimeter and millimeter wavelengths (mm) of high-redshift quasars trace the rest-frame far-infrared (FIR) emission from the dust of their host galaxies. Due to the negative K -correction, this provides the most efficient way to study the dusty star-forming interstellar medium (ISM) in the host galaxies. The early submillimeter/millimeter observations mainly focused on the samples of optically bright quasars, e.g. Max Planck Millimeter Bolometer array (MAMBO) survey on the Institute for Radio Astronomy in the Millimeter Range (IRAM) 30 m (Bertoldi et al. 2003a,b; Walter et al. 2003, 2009; Wang et al. 2007, 2010). The submillimeter/millimeter detection rate reached to 30% at mJy sensitivity and the FIR luminosities were 10^{12} - $10^{13} L_{\odot}$ (Beelen et al. 2006; Wang et al. 2007, 2008a). These detections indicated active star formation at rates of a few hundred to 1000 $M_{\odot} \text{ yr}^{-1}$, and dust masses of a few $10^8 M_{\odot}$ formed within 1 Gyr of the big bang.

Subsequently, Wang et al. (2011a) presented millimeter observations for an optically faint sample of 18 $z \sim 6$ quasars (5/18 detected at 250 GHz) with rest-frame 1450 Å magnitudes of $m_{1450} \geq 20.2$ mag, using the MAMBO-II on the IRAM 30 m telescope. Omont et al. (2013) also observed 20 $z \sim 6$ quasars with m_{1450} in the range of 19.63–24.15 mag using MAMBO, but only one quasar, CFHQS J142952+544717, was robustly detected. More recently, the technological improvement toward higher sensitivity has led to a large population of studies of $z \sim 6$ quasars targeting dust continuum performed with the Atacama Large Millimeter Array (ALMA), which detected the quasar host galaxies with FIR luminosities on the order of $10^{11} L_{\odot}$ and moderate star formation (Wang et al. 2013; Willott et al. 2013, 2015; Venemans et al. 2016, 2018; Decarli et al. 2018). In this work, we expand the submillimeter observations to a larger sample of quasars at the highest redshift, based on new observations from the Submillimetre Common-User Bolometer Array-2 (SCUBA2) camera on

the James Clerk Maxwell telescope (JCMT), to investigate the connection between FIR and AGN activities close to the epoch of cosmic reionization. In addition, compared to the previous works using the IRAM Northern Extended Millimeter Array (NOEMA) or ALMA, SCUBA2 provides a much larger field of view (15' in diameter), which allows us to study the environments of the quasars on megaparsec scales (Q. Li et al. 2020, in preparation).

Another peculiarity of $z \sim 6$ quasars regards their rest-frame ultraviolet (UV) emission lines. More than 30 quasars at $z \sim 6$ show weak $\text{Ly}\alpha$ emission. The $\text{Ly}\alpha$ lines from some of them even completely disappear in high-quality spectra; in addition, some objects show a heavy absorption feature (e.g., Bañados et al. 2015; Jiang et al. 2016). These quasars have $\text{Ly}\alpha + \text{Nv}$ rest-frame equivalent widths (EWs) of $< 15.4 \text{ \AA}$ (e.g., Fan et al. 2006; Bañados et al. 2014; Bañados et al. 2016; Jiang et al. 2016) and are called ‘weak line quasars’ (WLQs). Their EWs are much lower than the typical value of 62 \AA found with the normal SDSS quasars (Diamond-Stanic et al. 2009). In particular, Bañados et al. (2016) presented a sample 124 quasars from the Pan-STARRS1 (PS1) survey in the redshift range of $5.6 < z < 6.7$ and $M_{1450} < -25$ mag. They found 13.7% satisfy the weak-line quasar definition of Diamond-Stanic et al. (2009). Previous millimeter observations suggest that the high redshift mm-detected quasars tend to have weak UV emission line features (e.g., $z \sim 4$ optically bright quasars in Omont et al. 1996; $z \sim 6$ quasars in Bertoldi et al. 2003a; Wang et al. 2008b). However, the origin of this trend is not clear. By including our new JCMT data presented here, we could investigate this curious trend to the highest redshift covering a wide range of quasar luminosities from $2 \times 10^{13} L_{\odot}$ to $5 \times 10^{14} L_{\odot}$.

In this paper, we report our JCMT/SCUBA2 survey to study star formation in the host galaxies of 54 quasars at $5.6 < z < 6.9$, which expands the submillimeter observations to a large sample with a wide range of SMBH mass and quasar activity. This article is organized as follows. In Section 2, we introduce our sample selection criteria, observation procedures, and data reduction for the SCUBA2 survey; and refer briefly to the ancillary data we used. The observing results and the information of individual sources are presented in Section 3. In Section 4, we fit spectral energy distributions (SED) to calculate the FIR properties of $z \sim 6$ quasars and probe the relationship between far-infrared luminosity and AGN luminosity. In Section 5, we discuss the weak line features in this sample. Finally, we give a summary of the main results in Section 6. All magnitudes are

given in the AB system. Throughout this paper, we assume a flat cosmological model with $\Omega_\Lambda = 0.7$, $\Omega_M = 0.3$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (Spergel et al. 2007).

2. NEW SCUBA2 SURVEY OBSERVATIONS

2.1. Sample selection

We collect the $z > 5.6$ quasars discovered in recent years (e.g., Fan et al. 2006; Jiang et al. 2009; Willott et al. 2010; Bañados et al. 2014; Venemans et al. 2015b), and selected the objects that (i) have rest-frame 1450 Å absolute magnitudes of $M_{1450} < -25$ mag; (ii) are not included in previous single dish survey at 850 μm and 1.2 mm. (Wang et al. 2011a; Omont et al. 2013) The final sample of 54 quasars reported in this paper are shown in Figure 1 with the redshift range of 5.6–6.9 and 1450 Å magnitude range of $-27.6 \sim -25.0$ mag. They represent a quasar population at the highest redshift with quasar bolometric luminosities range of from $2 \times 10^{13} L_\odot$ to $5 \times 10^{14} L_\odot$ and SMBH masses from $5 \times 10^8 M_\odot$ to $1 \times 10^{10} M_\odot$.¹

2.2. Observations

The ‘SCUBA2 High rEdshift bRight quasaR surveY’ (hereafter SHERRY) was carried out by our team with the SCUBA2 camera on JCMT which is a 15m telescope located in Hawaii. The total observing time of SHERRY was 151.5 hrs including the overheads, scheduled flexibly (Program ID: M15BI055, M16AP013, M17AP062, M17BP034) from 2015 August to 2018 January. We used the constant velocity Daisy observing mode (‘CV DAISY’ mode) with the field of view of 15’ in diameter, which is designed for point/compact source observations. The resulting noise of the central circular region in radius of 5’5 is stable and of good quality (Chapin et al. 2013). SCUBA2 has two bands – 450 μm and 850 μm , which can be observed simultaneously. The main beam size of SCUBA2 is 7’’9 at 450 μm and 13’’0 at 850 μm . The pixel scale at 850 μm is 4’’/pixel while at 450 μm the sampling is 2’’/pixel.

The observations were carried out in grade 2 / grade 3 weather condition with the precipitable water vapor (PWV) in the range 0.83–2.58 mm, corresponding to the zenith atmospheric optical depth $0.05 < \tau_{225 \text{ GHz}} < 0.12$. We observed each target in four to six ~ 30 minutes scans with a total on-source time of 2–3 hrs per source, to reach the sensitivity of 1.2 mJy. The observed

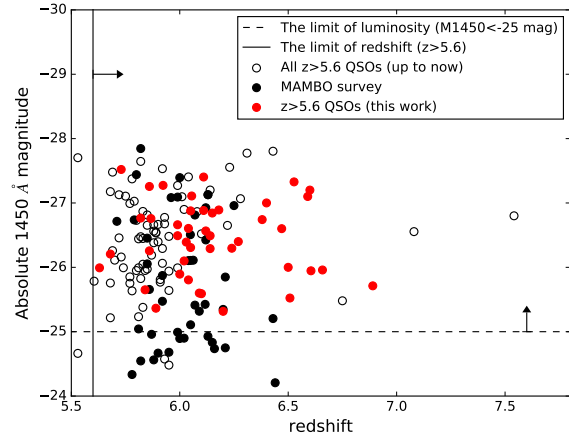


Figure 1. The sample selected from available quasar surveys before our observations. More than 250 quasars at $z > 5.6$ are known today (all circles). The red circles indicate the sources which belong to the sample in SHERRY, see Section 2.1.

calibration sources were taken before and after the target sources exposure, and selected from JCMT calibrator list (Dempsey et al. 2013), including Mars, Uranus, Neptune etc. The details of the observations are listed in Table 1.

2.3. Data reduction

The data were reduced using the STARLINK SCUBA2 science pipeline with the configuration file of ‘dimmconfig blank field.lis’, which is the recipe for processing maps containing faint compact sources (Chapin et al. 2013). Each complete observation was processed separately to produce an image, and calibrated with the flux conversion factors (FCFs) to mJy beam^{-1} . We adopted the default FCFs value of $537 \pm 24 \text{ Jy pW}^{-1} \text{ beam}^{-1}$ for 850 μm map and $491 \pm 67 \text{ Jy pW}^{-1} \text{ beam}^{-1}$ for 450 μm map (Dempsey et al. 2013). Then all the images for a given source were combined into a single co-added image using inverse-variance weighting. Using this recipe, the output map was further processed with a beam-match filtered of a 15’’ full-width-half-maximum (FWHM) Gaussian², then the S/N was taken to enhance point sources, which is suitable for quasars at $z \sim 6$. The resulting 850 and 450 μm maps have typical noise levels of 1.2 and 11.7 mJy beam^{-1} respectively,

¹ In this work, the AGN bolometric luminosities is estimated by the UV luminosities (1450 Å) with $L_{\text{bol}} = 4.2\nu L_{\nu, 1450}$ (Runnoe et al. 2012). To compared with previous work, we also recalculate the bolometric luminosities in other papers (e.g., Wang et al. 2011a; Omont et al. 2013) with the same conversion factor from Runnoe et al. (2012).

² In order to optimally find sources that are the size of the telescope beam, we used the PICARD recipe SCUBA2.MATCHED_FILTER with SMOOTH_FWHM = 15. It indicates that the background should be estimated by first smoothing the map and PSF with a 15-arcsec FWHM Gaussian. Thus, the beam size with a beam-match filtered is 15’’ at 850 μm .

which is comparable to previous observations of $z \sim 6$ quasars (e.g., ~ 0.5 mJy at 1.2 mm using MAMBO on the IRAM 30m telescope in Wang et al. 2007, assuming a gray body with dust temperature of 47K and emissivity index of 1.6 as in Beelen et al. 2006).

Then we used our clump-finding algorithm to find sources. The process identified pixels above 3σ and then checked whether they are within the region of one PSF. We considered $\geq 3\sigma$ signals with morphology consistent with the PSF as real detections, and adopt the peak value as the flux density of the source. Considering the beam size ($15''$ at $850 \mu\text{m}$ and $7''.9$ at $450 \mu\text{m}$), SCUBA2 peaks within the beam size from the optical quasar position were considered as the counterpart of the quasar.

2.4. Ancillary data

We also collected available optical to radio data from PS1, SDSS, Wide-field Infrared Survey Explorer (WISE), NOEMA, and ALMA surveys. The PS1 data are from the quasar discovery papers (e.g., Bañados et al. 2014;

Bañados et al. 2015, 2016; Venemans et al. 2015b); and WISE photometry data are from the ALLWISE Data Release (Cutri & et al. 2014), see Table 2. Twenty-one (21/54) objects have recent ALMA observations close to the [C II] $158 \mu\text{m}$ line frequency, around 250 GHz in observing frame (see the [C II] reference in Table 1).

The near-infrared (NIR) spectra of $z \sim 6$ quasars in our survey were published in the discovery papers (e.g. Morganson et al. 2012; Venemans et al. 2013, 2015b; Bañados et al. 2014; Jiang et al. 2015; Bañados et al. 2016; Jiang et al. 2016; Matsuoka et al. 2016; Wang et al. 2016, 2017). And we also include some unpublished sources (E. Bañados et al. 2020, in preparation; S. J. Warren et al. 2020, in preparation). These spectra have a wavelength range of 7,000–10,000 Å covering the Ly α and N v emission lines. The details are in Appendix B. We have corrected the Galactic extinction adopting the extinction curve presented in (Schlegel et al. 1998) for each spectra. We use these data to measure the equivalent width of Ly α and N v line and investigate the relationship between the sub-mm property and the weak emission-line feature.

Table 1. SCUBA2 survey observations

| source name | short name | RA | DEC | z | z_{method} | z_{ref} | UT date | T_{exp} | Weather |
|------------------------------|------------|-------------|-------------|----------------|---------------------|------------------|---|------------------|---------|
| (1) | (2) | (J2000) | (J2000) | (5) | (6) | (7) | (8) | (9) | (10) |
| | | | | | | | | (hour) | |
| SDSS J0008-0626 | J0008-0626 | 00:08:25.77 | -06:26:04.6 | 5.929±0.003 | O I | (1) | 2015-12-25; 2016-05-10 | 2.0 | band2 |
| PSO J002.3786+32.8702 | P002+32 | 00:09:30.88 | +32:52:12.9 | 6.10 | Template | (2) | 2017-09-14, 09-15 | 3.6 | band3 |
| PSO J007.0273+04.9571 | P007+04 | 00:28:06.56 | +04:57:25.7 | 6.0008±0.0004 | [C II] | (3) | 2015-12-17, 12-18 | 2.5 | band2 |
| SDSS J0100+2802 | J0100+2802 | 01:00:13.02 | +28:02:25.8 | 6.3258±0.0010 | [C II] | (4) | 2015-11-06, 11-15 | 2.0 | band2 |
| SDSS J0148+0600 | J0148+0600 | 01:48:37.64 | +06:00:20.0 | 5.923±0.003 | [C II] | (1) | 2015-11-04, 11-06, 11-15, 12-18 | 2.0 | band2 |
| PSO J036.5078+03.0498 | P036+03 | 02:26:01.87 | +03:02:59.4 | 6.54122±0.0018 | [C II] | (5) | 2017-08-30, 08-31, 09-12 | 3.6 | band3 |
| VIKJ0305-3150 | J0305-3150 | 03:05:16.91 | -31:50:55.9 | 6.6145±0.0001 | [C II] | (6) | 2017-08-02, 08-04, 08-14, 08-19, 08-20 | 6.2 | band3 |
| PSO J055.4244-00.8035 | P055-00 | 03:41:41.86 | -00:48:12.7 | 5.68 | template | (7) | 2015-10-23, 10-29 | 2.0 | band2 |
| PSO J056.7168-16.4769 | P056-16 | 03:46:52.04 | -16:28:36.8 | 5.99 | template | (2) | 2017-08-27, 08-28, 08-30 | 5.0 | band3 |
| PSO J060.5529+24.8567 | P060+24 | 04:02:12.69 | +24:51:24.4 | 6.18 | template | (2) | 2017-08-13, 08-24, 08-25, 08-26, 08-27, 08-31 | 3.6 | band3 |
| PSO J065.5041-19.4579 | P065-19 | 04:22:00.99 | -19:27:28.6 | 6.1247±0.0006 | [C II] | (3) | 2017-08-13, 08-19, 08-20, 08-24, 08-26, 08-27 | 4.5 | band3 |
| PSO J089.9394-15.5833 | P089-15 | 05:59:45.46 | -15:35:00.2 | 6.05 | template | (2) | 2017-08-04, 08-25, 08-26, 08-27, 08-31 | 4.2 | band3 |
| SDSS J0810+5105 | J0810+5105 | 08:10:54.32 | +51:05:40.0 | 5.82 | template | (2) | 2015-10-06, 10-07 | 2.5 | band2 |
| ULASJ0828+2633 (unpublished) | J0828+2633 | ... | ... | 6.05 | Other lines | (8) | 2017-02-17, 03-17, 03-20 | 2.0 | band2 |
| SDSS J0835+3217 | J0835+3217 | 08:35:25.76 | +32:17:52.6 | 5.89±0.03 | template | (9) | 2015-10-08, 10-23, 10-29 | 2.0 | band2 |
| SDSS J0839+0015 | J0839+0015 | 08:39:55.36 | +00:15:54.2 | 5.84±0.04 | template | (10) | 2016-03-19, 03-29 | 1.9 | band2 |
| SDSS J0842+1218 | J0842+1218 | 08:42:29.43 | +12:18:50.5 | 6.0763±0.0005 | [C II] | (3) | 2015-11-06, 11-07 | 2.0 | band2 |
| SDSS J0850+3246 | J0850+3246 | 08:50:48.25 | +32:46:47.9 | 5.867±0.007 | template | (1) | 2015-10-08, 10-22, 10-23 | 2.0 | band2 |
| PSO J135.3860+16.2518 | P135+16 | 09:01:32.65 | +16:15:06.8 | 5.63±0.05 | template | (7) | 2015-11-02, 11-03 | 2.0 | band2 |
| PSO J159.2257-02.5438 | P159-02 | 10:36:54.19 | -02:32:37.9 | 6.3809±0.0005 | [C II] | (3) | 2017-03-20, 03-21 | 2.0 | band2 |
| DELSJ104819.09-010940.21 | J1048-0109 | 10:48:19.08 | -01:09:40.3 | 6.6759±0.0005 | [C II] | (3) | 2017-03-22 | 2.0 | band2 |
| PSO J167.6415-13.4960 | P167-13 | 11:10:33.98 | -13:29:45.6 | 6.5148±0.0005 | [C II] | (3) | 2015-11-04, 11-06 | 2.0 | band2 |
| SDSS J1143+3808 | J1143+3808 | 11:43:38.33 | +38:08:28.7 | 5.80 | template | (2) | 2016-03-28, 03-29, 03-30 | 1.9 | band2 |
| SDSS J1148+0702 | J1148+0702 | 11:48:03.28 | +07:02:08.2 | 6.32 | Other lines | (8) | 2016-03-25, 03-28, 03-29 | 1.8 | band2 |
| HSCJ1152+0055 | J1152+0055 | 11:52:21.27 | +00:55:36.6 | 6.3637±0.0005 | [C II] | (11) | 2017-03-23, 03-25, 03-26, 03-27 | 2.5 | band2 |
| HSCJ1205-0000 | J1205-0000 | 12:05:05.09 | -00:00:27.9 | 6.73±0.02 | Mg II | (15) | 2017-03-20, 03-22 | 2.0 | band2 |
| SDSS J1207+0630 | J1207+0630 | 12:07:37.43 | +06:30:10.1 | 6.0366±0.0009 | [C II] | (3) | 2016-03-19 | 1.8 | band2 |
| PSO J183.1124+05.0926 | P183+05 | 12:12:26.98 | +05:05:33.5 | 6.4389±0.0004 | [C II] | (3) | 2017-03-26, 03-27 | 2.0 | band2 |
| PSO J183.2991-12.7676 | P183-12 | 12:13:11.81 | -12:46:03.5 | 5.86±0.02 | Other lines | (13) | 2016-03-18 | 2.1 | band2 |
| PSO J184.3389+01.5284 | P184+01 | 12:17:21.34 | +01:31:42.4 | 6.20 | template | (2) | 2018-03-27 | 2.0 | band2 |
| PSO J187.3050+04.3243 | P187+04 | 12:29:13.21 | +04:19:27.7 | 5.89±0.02 | N V | (13) | 2016-03-18 | 1.9 | band2 |
| SDSS J1243+2529 | J1243+2529 | 12:43:40.81 | +25:29:23.8 | 5.85 | Other line | (8) | 2016-03-25, 03-30, 04-01 | 1.8 | band2 |
| SDSS J1257+6349 | J1257+6349 | 12:57:57.47 | +63:49:37.2 | 6.02±0.03 | Ly α drop | (1) | 2016-02-24, 03-02, 03-05, 03-17 | 2.4 | band2 |
| PSO J210.4472+27.8263 | P210+27 | 14:01:47.34 | +27:49:35.0 | 6.14 | template | (2) | 2017-03-29, 03-30 | 2.0 | band2 |
| PSO J210.7277+40.4008 | P210+40 | 14:02:54.67 | +40:24:03.1 | 6.04 | template | (2) | 2017-02-15, 03-29, 03-30 | 3.0 | band2 |
| SDSS J1403+0902 | J1403+0902 | 14:03:19.13 | +09:02:50.9 | 5.86±0.03 | Ly α drop | (1) | 2016-02-09, 03-02, 03-25 | 1.8 | band2 |
| PSO J210.8722-12.0094 | P210-12 | 14:03:29.33 | -12:00:34.1 | 5.84±0.05 | template | (13) | 2016-02-21 | 2.1 | band2 |
| PSO J215.1514-16.0417 | P215-16 | 14:20:36.34 | -16:02:30.2 | 5.73±0.02 | O I | (14) | 2016-02-21, 03-17 | 2.2 | band2 |
| P215+26 (unpublished) | P215+26 | ... | ... | 6.27 | ... | (16) | 2017-04-15 | 2.0 | band2 |
| PSO J217.0891-16.0453 | P217-16 | 14:28:21.39 | -16:02:43.3 | 6.1498±0.0011 | [C II] | (3) | 2017-02-09, 02-17, 02-24 | 2.4 | band2 |
| PSO J217.9185-07.4120 | P217-07 | 14:31:40.45 | -07:24:43.4 | 6.14 | template | (2) | 2017-03-27, 03-29 | 2.0 | band2 |
| PSO J231.6576-20.8335 | P231-20 | 15:26:37.84 | -20:50:00.7 | 6.5864±0.0005 | [C II] | (3) | 2017-03-22, 03-23, 03-24 | 2.4 | band2 |
| PSO J239.7124-07.4026 | P239-07 | 15:58:50.99 | -07:24:09.5 | 6.11 | template | (2) | 2017-02-16, 02-17, 02-24, 04-15 | 2.5 | band2 |
| SDSS J1609+3041 | J1609+3041 | 16:09:37.27 | +30:41:47.6 | 6.16±0.03 | Mg II | (9) | 2016-02-04 | 1.8 | band2 |
| PSO J247.2970+24.1277 | P247+24 | 16:29:11.30 | +24:07:39.6 | 6.476 | Mg II | (15) | 2017-02-17, 02-18, 02-24 | 2.5 | band2 |
| PSO J261.0364+19.0286 | P261+19 | 17:24:08.74 | +19:01:43.1 | 6.44±0.05 | template | (15) | 2017-09-14, 09-15 | 1.8 | band3 |
| PSO J308.0416-21.2339 | P308-21 | 20:32:09.99 | -21:14:02.3 | 6.2341±0.0005 | [C II] | (3) | 2017-03-20, 03-21 | 2.4 | band2 |
| J2100-1715 | J2100-1715 | 21:00:54.61 | -17:15:22.5 | 6.0812±0.0005 | [C II] | (3) | 2017-08-14, 08-19, 08-27, 08-31 | 5.1 | band3 |
| PSO J323.1382+12.2986 | P323+12 | 21:32:33.18 | +12:17:55.2 | 6.5881±0.0003 | [C II] | (15) | 2017-08-18, 08-20, 08-27, 08-30 | 3.6 | band3 |
| PSO J333.9859+26.1081 | P333+26 | 22:15:56.63 | +26:06:29.4 | 6.03 | template | (2) | 2017-08-14, 08-19, 08-20 | 3.6 | band3 |
| PSO J338.2298+29.5089 | P338+29 | 22:32:55.15 | +29:30:32.2 | 6.666±0.004 | [C II] | (15) | 2015-12-17, 12-18, 12-25 | 2.0 | band2 |
| PSO J340.2041-18.6621 | P340-18 | 22:40:50.01 | -18:39:43.8 | 6.01 | template | (2) | 2017-08-31, 09-12, 09-14 | 5.1 | band3 |
| VIKJ2348-3054 | J2348-3054 | 23:48:33.33 | -30:54:10.2 | 6.9018±0.0007 | [C II] | (6) | 2017-07-05, 08-02, 08-04, 08-10, 08-14 | 6.8 | band3 |
| SWJ235632.44-062259.25 | P359-06 | 23:56:32.45 | -06:22:59.2 | 6.1722±0.0004 | [C II] | (3) | 2017-06-24, 08-28, 08-30 | 3.6 | band3 |

NOTE—

- a. (1) source name; (2) short source name; (3) and (4) RA and DEC in J2000; (5), (6) and (7) redshift, method used to estimate the redshift and references; (8), (9) and (10) JCMT observing date, exposure time and weather.
- b. Quasars sorted by right ascension. The reported coordinates are from the discovery papers.
- c. J0828+2633 and P215+26 are unpublished quasars (S. J. Warren et al. in prep.; Bañados et al. in prep.).
- d. Methods used to estimate the redshift: [C II]; optical lines; template fitting or Ly α . The redshift reported here is preferentially estimated from [C II] as it has less redshift uncertainties. If there's no radio observation, we report the redshift estimated from optical/NIR spectra. The reference papers list here: (1) Jiang et al. (2015); (2) Bañados et al. (2016); (3) Decarli et al. (2018); (4) Wang et al. (2016); (5) Bañados et al. (2015); (6) Venemans et al. (2016); (7) Bañados et al. (2015); (8) S. J. Warren et al. in prep.; (9) Jiang et al. (2016); (10) Venemans et al. (2015a); (11) Izumi et al. (2018); (12) Matsuoka et al. (2016); (13) Bañados et al. (2014); (14) Morganson et al. (2012); (15) Mazzucchelli et al. (2017); (16) Bañados et al. in prep.
- e. Weather Band 2 conditions are classified as dry, and translate to 850 μm transmissions of approximately 77% and 450 μm transmissions of approximately 19%. The precipitable water vapor (PWV) is in the range 0.83–1.58 mm, corresponding to the zenith atmospheric optical depth at 225 GHz $0.05 < \tau_{225 \text{ GHz}} < 0.08$. Weather Band 3 conditions translate to 850 μm transmissions of around 67% and 450 μm transmissions fall to approximately 7%. To reach the same sensitivity of 1.2 mJy at 850 μm , the exposure time was extended. The PWV at weather band 3 is 1.58–2.58 mm, corresponding to the zenith atmospheric optical depth of $0.08 < \tau_{225 \text{ GHz}} < 0.12$. (Dempsey et al. 2013)

Table 2. Quasar samples and SCUBA2 detections

| Source | M_{1450} | $S_{\nu, 850 \mu\text{m}}$ | offset | $S_{\nu, 450 \mu\text{m}}$ | offset | [C II] continuum | W1 | W2 | W3 | W4 | \hat{y}_{P1} | \hat{z}_{P1} | y_{P1} |
|------------------------------|------------|----------------------------|--------|----------------------------|--------|------------------|------------|------------|------------|------------|-----------------------|-----------------------|-----------------|
| (1) | (2) | (3) | (4) | (5) | (6) | (7) | (8) | (9) | (10) | (11) | (12) | (13) | (14) |
| | (mag) | (mJy) | ('') | (mJy) | ('') | (mJy) | (mag) | (mag) | (mag) | (mag) | (mag) | (mag) | (mag) |
| Detections: | | | | | | | | | | | | | |
| J0100+2802 | -29.2 | 4.09±1.13 | 4 | 2.52 ± 9.34 | ... | 1.35±0.25 | 17.16±0.03 | 16.98±0.03 | 16.89±0.21 | 15.60±0.44 | 20.76±0.04 | 18.61±0.01 | 17.62±0.01 |
| J0148+0600 | -27.3 | 5.28±1.19 | 0 | 1.74 ± 10.81 | ... | ... | 18.80±0.06 | 18.61±0.10 | 17.35±0.30 | 15.58±0.39 | 22.5 ± 0.09 | 19.45±0.01 | 19.37±0.04 |
| P036+03 | -27.3 | 5.47±1.01 | 0 | 5.02 ± 17.22 | ... | 2.5±0.5 | 19.42±0.08 | 19.47±0.18 | >17.20 | >15.56 | 23.54±0.24 | 21.44±0.12 | 19.26±0.03 |
| J0305-3150 | -25.9 | 8.43±1.08 | 4 | 0.62 ± 37.76 | ... | 3.29±0.10 | 20.38±0.15 | 20.09±0.24 | >18.02 | >15.60 | ... | ... | ... |
| P089-15 | -26.9 | 3.56±1.17 | 0 | 46.11 ± 31.67 | ... | ... | 20.38±0.19 | >20.25 | >17.97 | >15.85 | >23.17 | 19.66±0.03 | 19.85±0.09 |
| P135+16 | -26.0 | 5.16±1.34 | 0 | -5.16 ± 18.55 | ... | ... | 19.51±0.11 | 19.53±0.22 | 17.40±0.47 | >15.48 | 22.45±0.15 | 20.87±0.04 | 20.74±0.12 |
| J1048-0109 | -26.0 | 4.56±1.17 | 0 | 4.49 ± 13.11 | ... | 2.84±0.04 | 20.04±0.17 | 20.35±0.47 | >17.69 | >14.98 | >23.0 | >23.0 | 21.0 ± 0.15 |
| P183+05 | -27.0 | 9.03±1.30 | 5.7 | -7.87 ± 8.54 | ... | 4.47±0.02 | 19.66±0.13 | 19.79±0.33 | >17.68 | >15.54 | ... | 21.68±0.10 | 20.01±0.06 |
| P183-12 | -27.3 | 4.08±1.14 | 0 | 0.38 ± 9.64 | ... | ... | 18.98±0.07 | 19.18±0.16 | >17.59 | >15.53 | 22.14±0.13 | 19.47±0.02 | 19.23±0.03 |
| P215-16 | -27.6 | 16.85±1.10 | 0 | 26.93±7.78 | 2.0 | ... | 18.35±0.05 | 18.52±0.08 | 16.32±0.11 | >15.12 | 21.48±0.05 | 19.08±0.02 | 19.14±0.03 |
| P217-07 | -26.4 | 6.03±1.17 | 5.7 | 3.32 ± 11.64 | ... | ... | 19.88±0.12 | 19.81±0.25 | >17.30 | >15.45 | >23.79 | 21.1 ± 0.08 | 20.44±0.08 |
| P231-20 | -27.2 | 7.99±1.22 | 0 | 80.31±19.97 | 2.8 | 4.41±0.16 | 19.91±0.15 | 19.96±0.35 | >17.41 | >15.56 | ... | >22.77 | 20.14±0.08 |
| J1609+3041 | -26.6 | 4.09±1.17 | 4 | -0.59 ± 7.36 | ... | ... | 20.22±0.14 | 20.42±0.35 | >17.64 | >15.88 | 23.55±0.27 | 20.98±0.07 | 20.43±0.09 |
| P308-21 | -26.4 | 4.23±1.09 | 4 | -6.52 ± 17.11 | ... | 1.34±0.21 | 19.20±0.09 | 18.78±0.13 | >17.62 | >14.77 | 23.58±0.27 | 21.12±0.08 | 20.49±0.11 |
| P333+26 | -26.4 | 3.83±1.04 | 4 | -32.49 ± 18.71 | ... | ... | 19.83±0.11 | >20.12 | >17.98 | >15.85 | >23.53 | 20.91±0.09 | 20.38±0.1 |
| J2348-3054 | -25.8 | 5.88±1.06 | 0.0 | 9.48 ± 55.42 | ... | 1.92±0.14 | 20.37±0.17 | >20.07 | >17.31 | 15.76±0.49 | ... | ... | ... |
| Tentative detections: | | | | | | | | | | | | | |
| P187+04 | -25.4 | 3.81±1.25 | 12.6 | -19.49 ± 11.46 | ... | ... | 19.56±0.12 | >19.96 | >17.03 | >15.40 | 23.44±0.26 | 20.92±0.04 | 21.11±0.13 |
| J1157+6349 | -26.6 | 3.34±1.08 | 8 | 10.95 ± 8.00 | ... | ... | 19.61±0.08 | 19.86±0.19 | >17.69 | >15.38 | 22.85±0.21 | 20.8 ± 0.07 | 20.5 ± 0.14 |
| P210-12 | -25.7 | 3.56±1.16 | 8 | 4.77 ± 9.78 | ... | ... | 18.92±0.06 | 19.49±0.20 | >17.45 | >15.78 | >23.21 | 21.09±0.07 | 20.89±0.13 |
| P247+24 | -26.6 | 6.76±1.13 | 8.9 | 9.18 ± 13.82 | ... | ... | 19.49±0.08 | 19.41±0.14 | 17.46±0.40 | >15.19 | ... | >22.77 | 20.04±0.07 |
| Non-detections: | | | | | | | | | | | | | |
| J0008-0626 | -26.1 | 2.15±1.19 | ... | 8.02±8.99 | ... | ... | 19.51±0.11 | 19.02±0.14 | >17.07 | >15.38 | 22.68±0.13 | 20.25±0.03 | 20.42±0.10 |
| P002+32 | -25.6 | 0.00±1.14 | ... | -28.77±28.47 | ... | ... | 20.42±0.21 | >20.32 | >17.98 | >15.72 | >23.82 | 21.18±0.06 | 21.15±0.15 |
| P007+04 | -26.6 | 0.03±1.07 | ... | -13.38±10.49 | ... | 2.07±0.04 | 19.97±0.16 | 19.75±0.29 | >17.43 | >14.94 | 22.84±0.16 | 20.56±0.05 | 20.12±0.07 |
| P055-00 | -26.2 | 0.37±1.24 | ... | 14.61±12.34 | ... | ... | ... | ... | ... | ... | 22.27±0.13 | 20.19±0.04 | 20.29±0.09 |
| P056-16 | -26.7 | -0.62±0.89 | ... | -8.49±18.58 | ... | ... | 18.32±0.04 | 19.08±0.12 | >17.86 | >15.44 | 22.99±0.28 | 20.0 ± 0.04 | 20.05±0.08 |
| P060+24 | -26.9 | 0.84±1.19 | ... | 15.36±25.34 | ... | ... | 19.17±0.09 | 19.41±0.21 | >17.19 | >15.46 | 23.01±0.3 | 20.18±0.03 | 19.87±0.06 |
| P065-19 | -26.6 | 2.50±0.98 | ... | -29.34±23.35 | ... | 0.46±0.05 | 20.09±0.14 | 20.19±0.30 | >17.59 | >15.30 | 23.47±0.24 | 19.79±0.03 | 20.19±0.08 |
| J0810+5105 | -26.8 | 0.75±1.15 | ... | 10.16±14.29 | ... | ... | 19.64±0.10 | 19.20±0.14 | >17.54 | >15.61 | 22.29±0.1 | 19.7 ± 0.02 | 19.89±0.06 |
| J0828+2633 | -26.4 | 2.56±1.18 | ... | 1.85±14.12 | ... | ... | 16.22±0.03 | 16.86±0.04 | >17.82 | >15.44 | >23.44 | 20.72±0.06 | 20.42±0.1 |
| J0835+3217 | -25.8 | -0.32±1.24 | ... | -17.36±15.76 | ... | ... | 20.31±0.22 | 20.47±0.54 | >17.67 | >15.31 | ... | ... | ... |
| J0839+0015 | -25.4 | 1.68±1.18 | ... | -3.93±9.12 | ... | ... | 20.15±0.18 | >19.90 | 17.64±0.50 | >15.58 | >23.59 | 21.18±0.09 | >21.68 |
| J0842+1218 | -26.9 | 2.30±1.25 | ... | -2.05±13.11 | ... | 0.65±0.06 | 18.92±0.07 | 19.31±0.17 | 17.53±0.43 | >15.64 | >23.43 | 19.83±0.03 | 19.88±0.06 |
| J0850+3246 | -26.8 | 2.09±1.21 | ... | 18.64±13.50 | ... | ... | 19.43±0.09 | 18.98±0.15 | >17.13 | >15.21 | 22.48±0.15 | 20.17±0.03 | 20.0 ± 0.05 |
| P159-02 | -26.8 | 0.85±1.27 | ... | -23.56±13.80 | ... | 0.65±0.03 | 19.44±0.09 | 19.04±0.13 | 17.65±0.51 | >15.52 | >23.62 | 20.46±0.04 | 19.88±0.06 |
| P167-13 | -25.6 | 2.39±1.18 | ... | 2.06±9.46 | ... | 0.87±0.05 | 20.48±0.21 | >20.66 | >17.70 | >15.50 | >23.45 | >22.86 | 20.48±0.11 |
| J1143+3808 | -26.8 | 1.92±1.32 | ... | -1.60±15.65 | ... | ... | 19.63±0.10 | 19.37±0.16 | >17.97 | >15.58 | 22.43±0.08 | 20.11±0.03 | 20.02±0.05 |
| J1148+0702 | -26.4 | -1.48±1.30 | ... | -8.78±13.19 | ... | 0.41±0.05 | 19.09±0.08 | 18.82±0.13 | >17.26 | >14.86 | 23.28±0.26 | 20.99±0.07 | 20.37±0.11 |
| J1152+0055 | -25.1 | -0.82±1.05 | ... | -4.61±9.15 | ... | 0.1890±0.32 | 20.10±0.19 | 20.32±0.48 | >17.58 | >14.95 | ... | ... | ... |
| J1205-0000 | -25.0 | 2.05±1.21 | ... | 0.833±13.45 | ... | 0.833±0.176 | 19.98±0.15 | 19.65±0.23 | >17.73 | >15.44 | ... | ... | ... |
| J1207+0630 | -26.6 | 2.31±1.22 | ... | 9.57±9.21 | ... | 0.50±0.06 | 19.69±0.13 | 19.43±0.21 | >17.21 | >14.84 | 22.98±0.17 | 20.44±0.03 | 20.15±0.17 |
| P184+01 | -25.4 | -1.92±1.11 | ... | -7.15±9.74 | ... | ... | 20.28±0.21 | >19.60 | >17.32 | >14.81 | >23.74 | 21.2 ± 0.07 | 21.46±0.2 |

Table 2 continued

Table 2 (continued)

| Source | M_{1450} | $S_{\nu, 850 \mu\text{m}}$ | offset | $S_{\nu, 450 \mu\text{m}}$ | offset | [C II] continuum | W1 | W2 | W3 | W4 | $^i\text{P1}$ | $^z\text{P1}$ | $^y\text{P1}$ |
|------------|------------|----------------------------|----------|----------------------------|----------|-------------------|------------------|------------------|------------------|------------------|------------------|------------------|------------------|
| | (mag) | (mJy) | ($''$) | (mJy) | ($''$) | (mJy) | (mag) | (mag) | (mag) | (mag) | (mag) | (mag) | (mag) |
| (1) | (2) | (3) | (4) | (5) | (6) | (7) | (8) | (9) | (10) | (11) | (12) | (13) | (14) |
| J1243+2529 | -26.2 | 0.43 ± 1.34 | ... | -1.14 ± 14.04 | ... | ... | 19.40 ± 0.09 | 18.86 ± 0.11 | > 17.73 | > 15.43 | 23.43 ± 0.26 | 20.24 ± 0.03 | 20.63 ± 0.08 |
| P210+27 | -26.5 | -1.11 ± 1.17 | ... | -1.75 ± 10.14 | ... | ... | 20.26 ± 0.16 | 20.32 ± 0.35 | > 17.51 | > 15.61 | > 23.75 | 21.18 ± 0.06 | 20.27 ± 0.08 |
| P210+40 | -25.9 | -1.20 ± 1.24 | ... | 15.26 ± 15.01 | ... | ... | 19.28 ± 0.07 | 19.56 ± 0.18 | > 17.65 | > 15.43 | > 23.86 | 20.87 ± 0.05 | 20.92 ± 0.12 |
| J1403+0902 | -26.3 | -0.55 ± 1.39 | ... | 7.01 ± 14.21 | ... | ... | 19.98 ± 0.12 | 20.09 ± 0.29 | > 17.44 | 15.67 ± 0.42 | 23.02 ± 0.16 | 20.31 ± 0.03 | 20.31 ± 0.09 |
| P215+26 | -26.4 | 2.11 ± 1.19 | ... | -9.35 ± 10.76 | ... | ... | 19.55 ± 0.08 | 20.56 ± 0.37 | 17.36 ± 0.28 | > 15.85 | ... | ... | ... |
| P217-16 | -26.9 | 1.48 ± 1.21 | ... | -11.24 ± 18.40 | ... | 0.37 ± 0.06 | 17.11 ± 0.03 | 17.77 ± 0.05 | > 17.53 | > 15.81 | 23.18 ± 0.29 | 20.46 ± 0.04 | 19.87 ± 0.06 |
| P239-07 | -27.5 | 2.56 ± 1.15 | ... | 9.30 ± 11.67 | ... | ... | 16.48 ± 0.14 | 17.17 ± 0.04 | > 17.31 | > 15.45 | 22.93 ± 0.24 | 19.78 ± 0.03 | 19.34 ± 0.04 |
| P261+19 | -26.0 | 1.94 ± 1.62 | ... | -40.85 ± 51.07 | ... | ... | 19.58 ± 0.09 | 20.51 ± 0.40 | > 17.74 | > 15.48 | ... | > 22.92 | 20.98 ± 0.13 |
| J2100-1715 | -25.6 | 0.46 ± 0.95 | ... | 7.28 ± 24.56 | ... | 0.52 ± 0.06 | 18.63 ± 0.06 | 19.23 ± 0.17 | > 17.49 | > 15.19 | ... | ... | ... |
| P323+12 | -27.1 | 0.56 ± 1.03 | ... | 2.91 ± 18.30 | ... | 0.47 ± 0.146 | 19.06 ± 0.07 | 18.97 ± 0.12 | > 17.49 | > 15.57 | ... | 21.56 ± 0.10 | 19.28 ± 0.03 |
| P338+29 | -26.0 | 0.10 ± 1.31 | ... | -13.71 ± 14.86 | ... | 0.972 ± 0.215 | 20.51 ± 0.21 | > 20.04 | > 17.22 | > 15.56 | > 23.29 | > 22.5 | 20.23 ± 0.1 |
| P340-18 | -26.4 | -1.05 ± 1.02 | ... | -3.41 ± 31.73 | ... | 0.13 ± 0.05 | 19.26 ± 0.09 | 18.87 ± 0.13 | > 17.51 | > 15.30 | 23.34 ± 0.29 | 20.14 ± 0.03 | 20.35 ± 0.1 |
| P359-06 | -26.8 | 1.52 ± 1.10 | ... | -3.22 ± 25.80 | ... | 0.87 ± 0.09 | 19.27 ± 0.10 | 19.72 ± 0.41 | > 17.36 | > 15.35 | 23.02 ± 0.21 | 19.97 ± 0.03 | 20.03 ± 0.06 |

NOTE—

- (1) source name; (2) rest-frame 1450Å absolute magnitudes; (3)-(6) 850μm and 450μm flux densities and the offsets between the quasar optical position and the sub-mm detected brightest pixel; (7) [C II] continuum flux density; (8) - (11) WISE magnitudes from the ALLWISE source catalog; (12) - (14) the dereddened PSI magnitudes (Bañados et al. 2016).
- The detections of 450 or 850 μm band are marked in boldface.
- The offsets represent the distance between SCUBA2 detected position and its quasar optical position.
- Other bands data is from WISE, Pan STARRS1, and the radio papers (see Table 1); and all magnitudes are given in the AB system.
- For $> 3\sigma$ objects that are more than half beam away from the quasar optical positions, we list them as 'tentative detections'.

3. RESULTS AND ANALYSIS

3.1. Source detection

Following the source detection criteria described in the previous section, twenty quasars (20/54) are detected at the $\geq 3\sigma$ level or above (> 4 mJy) at $850\ \mu\text{m}$ (shown in Figure 2), and two are detected at $450\ \mu\text{m}$. As the sources are un-resolved, for detections we adopt the peak value close to the quasar position as the total flux density; and for non-detections, we adopt the pixel value at the optical quasar position. The results are presented in Table 2, including the source name, M_{1450} , 450 and $850\ \mu\text{m}$ flux densities, and the offset between the SCUBA2 peak and the optical quasar position. We also include the data at other wavelengths from the literature. Their NIR spectral information are listed in Appendix B. The EW value quoted here are literature values. The details for the individual detected sources are described as follows:

SDSS J010013.02+280225.8 (hereafter J0100+2802) at $z = 6.30$ is a weak line quasar discovered by Wu et al. (2015) with EW (Ly α + N v) $\sim 10\ \text{\AA}$ from the LBT spectrum. It is the only known quasar with a bolometric luminosity higher than $10^{48}\ \text{erg s}^{-1}$ and a black-hole mass larger than $5 \times 10^9 M_{\odot}$ at $z \geq 6$. Wang et al. (2016) report that its 260 GHz flux density is 1.35 ± 0.25 mJy using Plateau de Bure interferometer (PdBI). This source is detected in our SCUBA2 $850\ \mu\text{m}$ observation at $\sim 3.6\sigma$ level, with a flux density of $S_{\nu,850\ \mu\text{m}} = 4.09 \pm 1.13$ mJy, $S_{\nu,450\ \mu\text{m}} < 30.52$ mJy, $4.0''$ away from the optical position.

SDSS J014837.64+060020.0 (hereafter J0148+0600) at $z = 5.923 \pm 0.003$ is a low-ionization BAL (LoBAL) quasar discovered by S. Warren et al. (2015, in preparation) with EW(Ly α) $> 87\ \text{\AA}$. Our SCUBA2 $850\ \mu\text{m}$ observation shows $\sim 4.4\sigma$ detection with a flux density of $S_{\nu,850\ \mu\text{m}} = 5.27 \pm 1.19$ mJy, consistent with the optical position. The source is undetected at $450\ \mu\text{m}$, with a 3σ upper limit of < 34.17 mJy.

PSO J036.5078+03.0498 (hereafter P036+03) is one of the most luminous objects ($M_{1450} = -27.36 \pm 0.03$ mag) known at $z > 6$ discovered by Venemans et al. (2015b) with a redshift $z_{\text{MgII}} = 6.527 \pm 0.002$, and is also detected in the UKIDSS (Warren et al. 2007) and WISE catalogs (Cutri & et al. 2014). Later, Bañados et al. (2015) reported a strong detection of the [C II] line ($L_{[\text{CII}]} = 5.8 \pm 0.7 \times 10^9 L_{\odot}$) in the host galaxy of this source using the IRAM NOEMA millimeter interferometer, yielding a redshift of $z_{[\text{CII}]} = 6.54122 \pm 0.0018$. Our SCUBA2 observation shows $\sim 5.4\sigma$ detection at $850\ \mu\text{m}$ ($S_{\nu,850\ \mu\text{m}} = 5.47 \pm 1.01$ mJy) at the optical position.

VIKING J030516.92-315056.0 (hereafter J0305-3150) is a luminous quasar ($M_{1450} = -25.96 \pm 0.06$ mag) discovered by Venemans et al. (2013) using the Magellan/FIRE telescope. The precise redshift is $z = 6.6145 \pm 0.0001$ given by [C II] line (Venemans et al. 2016). It is detected $\sim 7.8\sigma$ with its flux density of $S_{\nu,850\ \mu\text{m}} = 8.43 \pm 1.08$ mJy in our SCUBA2 survey, $4''$ away from the optical position.

PSO J089.9394-15.5833 (hereafter J089-15) is a PS1 discovered quasar and spectroscopically confirmed using Low-Resolution Imaging Spectrometer at the Keck I (Bañados et al. 2016). This source is detected $\sim 3.0\sigma$ level at the $850\ \mu\text{m}$ band, with a flux density of $S_{\nu,850\ \mu\text{m}} = 5.56 \pm 1.17$ mJy at the optical position.

PSO J135.3860+16.2518 (hereafter P135+16) is a $z = 5.63 \pm 0.05$ radio loud quasar discovered by Bañados et al. (2015) with $R = 91.4 \pm 8.8$. It is detected at 1.4 GHz by the VLA with $S_{1.4\ \text{GHz,peak}} = 3.04 \pm 0.15$ mJy. We detect a $\sim 3.8\sigma$ peak ($S_{\nu,850\ \mu\text{m}} = 5.15 \pm 1.34$ mJy, $S_{\nu,450\ \mu\text{m}} < 50.50$ mJy) on the SCUBA2 map at the position of the optical quasar.

DELS J104819.09-010940.21 (hereafter J1048-0109) at $z = 6.63$ is the first $z > 6.5$ quasar discovered using DECaLS imaging and identified with Magellan/FIRE (Wang et al. 2017). It was also independently discovered using VIKING imaging by Venemans et al. in prep.. Venemans et al. (2018) detected this source using ALMA with $S_{\text{rest},1790\ \text{GHz}} = 2.722 \pm 0.094$ mJy and $S_{\text{rest},1900\ \text{GHz}} = 3.110 \pm 0.120$ mJy. Our $850\ \mu\text{m}$ SCUBA2 map shows $\sim 3.9\sigma$ detection ($S_{\nu,850\ \mu\text{m}} = 4.56 \pm 1.17$ mJy) at the quasar optical position.

PSO J183.1124+05.0926 (hereafter P183+05) is a $z = 6.4$ quasar discovered by Bañados et al. (2016). This object is a metal-poor proximate Damped Lyman Alpha system (DLA) (Bañados et al. 2019). Decarli et al. (2018) reported a strong detection of the [C II] line ($\log L_{[\text{CII}]} = 9.83 L_{\odot}$) with dust continuum of 4.47 ± 0.02 mJy at 250 GHz using ALMA. It is detected $\sim 6.9\sigma$ in $850\ \mu\text{m}$ map with its flux density of $S_{\nu,850\ \mu\text{m}} = 9.03 \pm 1.30$ mJy, $6''$ away from the optical position.

PSO J183.2991-12.7676 (hereafter P183-12) is a weak emission line quasar at $z = 5.86 \pm 0.02$ with the EW (Ly α + N v) of $11.8\ \text{\AA}$ (Bañados et al. 2014). It does not show any detectable emission line and it seems that Ly α is almost completely absorbed. Our SCUBA2 observation shows a $\sim 3.6\sigma$ peak detection at $850\ \mu\text{m}$ with a flux density of $S_{\nu,850\ \mu\text{m}} = 4.08 \pm 1.14$ mJy, and a $450\ \mu\text{m}$ upper limit of 29.30 mJy, at the optical position of the quasar.

PSO J187.3050+04.3243 (hereafter P187+04) is confirmed spectroscopically using the FOcal Reducer/ low

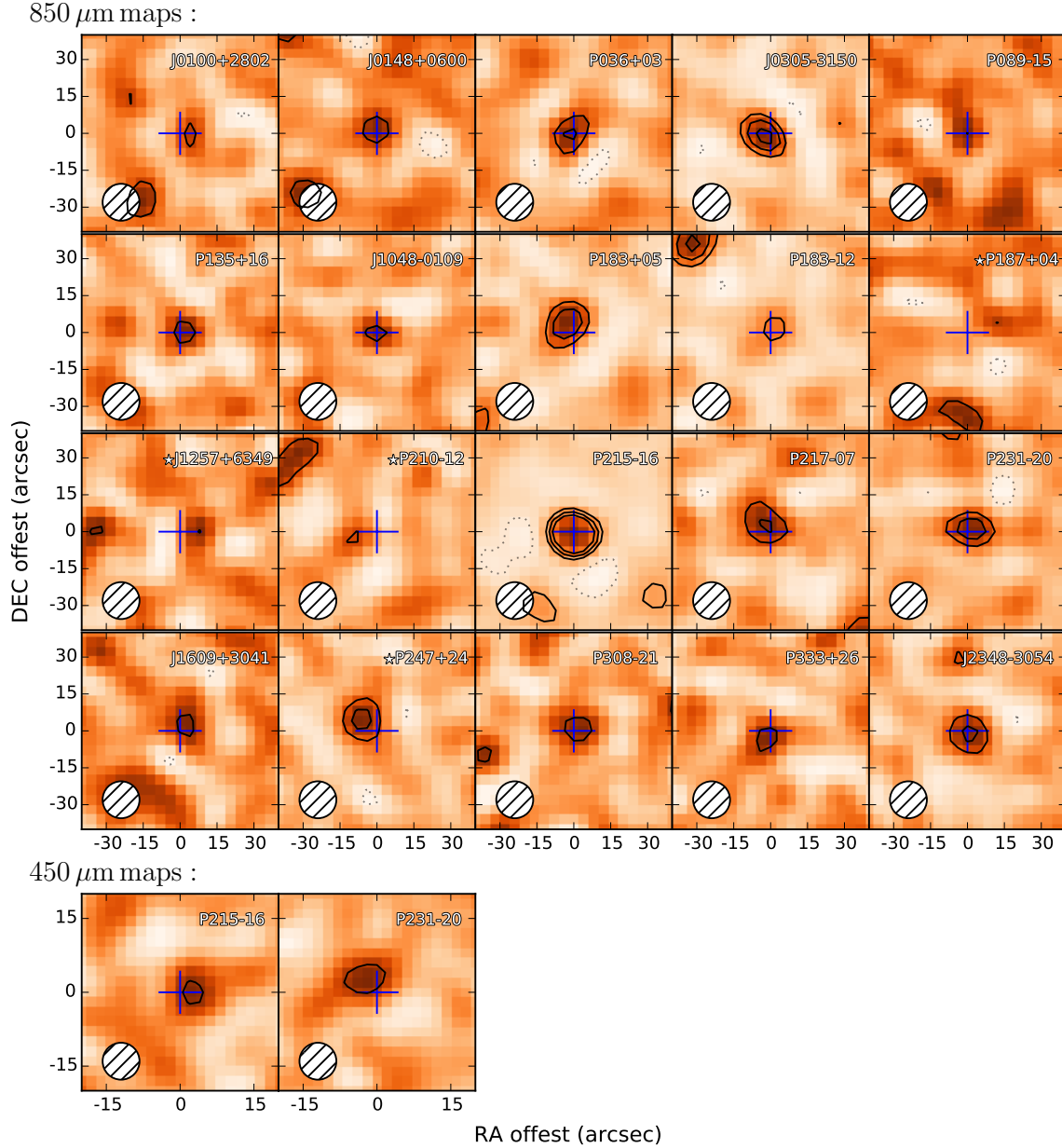


Figure 2. SCUBA2 850 μm and 450 μm detected quasars images at $z \sim 6$. The beam is about $15''$ indicated by the circle in 850 μm map; while $7''.9$ in 450 μm map. The blue cross marks the quasar optical position. The colour scale in mJy beam^{-1} goes from white to orange, with orange areas indicating higher flux. We scaled them to the same flux scales. The dotted and solid lines indicate the contour levels of $[-3]$ and $[+3, 5, 7\sigma]$ in each image. The tentative detection is marked with a star.

dispersion Spectrograph 2 (FOR2) at the Very Large Telescope (VLT) by Bañados et al. (2014). The FOR2 discovery spectrum shows bright and narrow $\text{Ly}\alpha$ and Nv emission lines. There is a tentative detection in the 850 μm map ($\sim 3.0\sigma$, $S_{\nu,850\mu\text{m}} = 3.81 \pm 1.25 \text{ mJy}$), $12.6''$ away from the optical position.

SDSSJ125757.47+634937.2 (hereafter J1257+6349) at $z = 6.02$ is a quasar with the $\text{EW}(\text{Ly}\alpha + \text{Nv})$ of 18 \AA discovered by Jiang et al. (2015). Its redshift was measured by the sharp flux drop at the $\text{Ly}\alpha$ line. SCUBA2 850 μm map shows a tentative detection of $\sim 3.1\sigma$ with a flux density of $S_{\nu,850\mu\text{m}} = 3.34 \pm 1.08 \text{ mJy}$, $8''$ away from the position of the optical quasar.

PSO J210.8722–12.0094 (hereafter P210–12) at $z = 5.84$ is the faintest of the PS1 quasar sample ($z_p = 21.15 \pm 0.08$ mag) discovered by Bañados et al. (2014). The VLT/FORS2 discovery spectrum shows a bright continuum and devoid of bright emission lines with the EW (Ly α + N v) of 10.7 \AA . The redshift is estimated by matching the continuum to the composite spectra from Vanden Berk et al. (2001) and Fan et al. (2006) (Bañados et al. 2014). It has a tentative detection in $850 \mu\text{m}$ map ($\sim 3.1\sigma$, $S_{\nu,850\mu\text{m}} = 3.56 \pm 1.16$ mJy), $8''$ away from the optical position.

PSO J215.1512–16.0417 (hereafter P215–16) is a $z = 5.73$ broad absorption line (BAL) quasar discovered by Morganson et al. (2012) with EW (Ly α + N v) of $107 \pm 83 \text{ \AA}$. It has a bolometric luminosity of $3.8 \times 10^{47} \text{ erg s}^{-1}$, and a black hole mass of $6.9 \times 10^9 M_\odot$ with the LBT Near-Infrared Spectroscopic. Our SCUBA2 observation detected a bright $\sim 15.3\sigma$ peak with a flux density for this source of $S_{\nu,850\mu\text{m}} = 16.85 \pm 1.10$ mJy at the position of the optical quasar; and $S_{\nu,450\mu\text{m}} = 26.93 \pm 7.78$ mJy. It is the brightest sub-mm source in SHERRY sample.

PSO J217.9185–07.4120 (hereafter P217–07) at $z = 6.14$ is a PS1 quasar and spectroscopically confirmed using the Low-Dispersion Survey Spectrograph (LDSS3) on Magellan (Bañados et al. 2016). It is detected by our SCUBA2 $850 \mu\text{m}$ observation at the $\sim 5.2\sigma$ level, with a flux density of $S_{\nu,850\mu\text{m}} = 6.03 \pm 1.17$ mJy, $5.7''$ away from the optical position.

PSO J231.6576–20.8335 (hereafter P231–20) is a $z = 6.6$ quasar at discovered by Mazzucchelli et al. (2017). It has bright detections in both $850 \mu\text{m}$ and $450 \mu\text{m}$ band. The flux densities are $S_{\nu,850\mu\text{m}} = 7.99 \pm 1.22$ mJy ($\sim 6.5\sigma$, right at the optical position) and $S_{\nu,450\mu\text{m}} = 80.31 \pm 19.97$ mJy ($\sim 4.0\sigma$, $2.8''$ away from the optical position). P231–20 has the dust continuum detected using ALMA of $F_{250\text{GHz}} = 4.41 \pm 0.16$ mJy with its companion of $F_{250\text{GHz}} = 1.73 \pm 0.16$ mJy in a SCUBA2 beam (Decarli et al. 2017). We estimate SCUBA2 flux density has 72% from the quasar using the ALMA continuum flux ratio between the quasar and its companion.

SDSS J1609+3041 (hereafter J1609+3041) is a $z \sim 6.16$ quasar discovered by Jiang et al. (2016) and independently discovered by UKIDSS (S. Warren et al. 2016, in preparation). It has tentative 1.4 GHz detections of $484 \pm 137 \mu\text{Jy}$ (S/N of 3.5) with radio loudness of $R = 28.3 \pm 8.6$ (Bañados et al. 2015). This source is detected at $\sim 3.5\sigma$ level by our SCUBA2 $850 \mu\text{m}$ observation, with a flux density for this source of $S_{\nu,850\mu\text{m}} = 4.09 \pm 1.17$ mJy, $S_{\nu,450\mu\text{m}} < 21.49$ mJy, $4''$ away from the position of the optical quasar.

PSO J247.2970+24.1277 (hereafter P247+24) at $z_{\text{MgII}} = 6.476 \pm 0.004$ is discovered from the PS1 survey and confirmed with VLT/FORS2 and Magellan/FIRE (Mazzucchelli et al. 2017). It is also detected at $850 \mu\text{m}$ band of $\sim 5.9\sigma$, with a flux density for this source of $S_{\nu,850\mu\text{m}} = 6.76 \pm 1.13$ mJy, $8.9''$ away from the position of the optical quasar.

PSO J308.0416–21.2339 (hereafter P308–21) at $z = 6.24$ is a PS1-discovered quasar and confirmed with VLT/FORS2 (Bañados et al. 2016). Our SCUBA2 survey at $850 \mu\text{m}$ band shows a $\sim 3.9\sigma$ detection, with a flux density for this source of $S_{\nu,850\mu\text{m}} = 4.23 \pm 1.09$ mJy, $4''$ away from the quasar optical position. The dust continuum using ALMA of P308–21 is $F_{250\text{GHz}} = 1.34 \pm 0.21$ mJy with its companion of $F_{250\text{GHz}} = 0.19 \pm 0.06$ mJy in a SCUBA2 beam (Decarli et al. 2017).

PSO J333.9859+26.1081 (hereafter P333+26) at $z = 6.03$ is a PS1-discovered quasar and confirmed with Keck/LRIS (Bañados et al. 2016). It is also detected in ALLWISE catalog. It is detected in SCUBA2 $850 \mu\text{m}$ map of $S_{\nu,850\mu\text{m}} = 3.83 \pm 1.04$ mJy ($\sim 3.7\sigma$), $4''$ away from the quasar optical position.

VIKING J234833.34–305410.0 (hereafter J2348–3054) is discovered using VLT/FORS2, which shows an absorption shortward of Ly α (Venemans et al. 2013). Later, it is confirmed as a BAL quasar using the VLT/X-Shooter spectrum (Venemans et al. 2013). The redshift is $z = 6.886 \pm 0.009$ measured from the Mg II line. Venemans et al. (2016) reported it has a [C II] and continuum detection using ALMA with $S_{\text{obs},1\text{mm}} = 1.92 \pm 0.14$ mJy. It is detected in the SCUBA2 $850 \mu\text{m}$ map of $\sim 5.5\sigma$ ($S_{\nu,850\mu\text{m}} = 5.88 \pm 1.06$ mJy) at the quasar optical position.

For $> 3\sigma$ objects that are more than half beam away from the quasar positions, we list them as tentative detections (e.g. P187+04, J1257+6349, P210–12). P247+24 has a good S/N of 6 but $8''9$ away from the optical position, we cannot rule out if there are some companions, thus we also list it as tentative detection. Detections at $\geq 3\sigma$ were also obtained in another three images at $850 \mu\text{m}$ (P007+04, P184+01, P210+40). However, the peaks are $14''–17''$ away from their optical quasar positions. Thus we do not consider them as the sub-mm counterparts of the quasar hosts. For the non-detections, we list the measurements at the on-source pixel in Table 2.

3.2. The average sub-mm properties of $z \sim 6$ quasar

We constructed three stacked averages: (i) the whole sample, (ii) the $850 \mu\text{m}$ detections, and (iii) the $850 \mu\text{m}$ non-detections. Here we included all the tentative detections in the $850 \mu\text{m}$ detection subsample. We stacked the

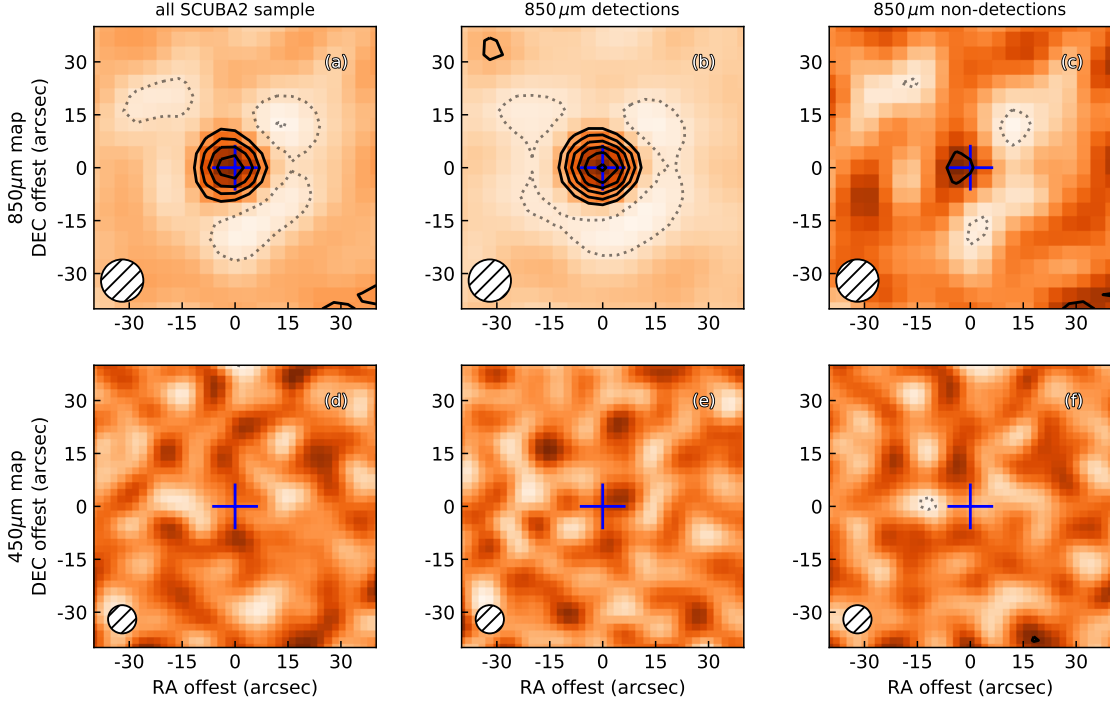


Figure 3. 80×80 arcsec² SCUBA2 850 μm and 450 μm stacked maps of (a, d) all the SHERRY sample, (b, e) 850 μm detections and (c, f) 850 μm non-detections. The blue cross denotes the quasar optical position. The white circle shows the SCUBA2 beam size at 850 μm of $15''$ and at 450 μm of $7''.9$. The colour scale in mJy beam^{-1} goes from white to orange with the contour levels of $[-6, -3]$ and $[+3, 6, 9\dots] \times \sigma$ (dotted and solid lines). The stacking procedure used to produce this image is described in Section 3.2.

Table 3. The median and weighted average parameters of $z \sim 6$ quasars in SHERRY

| Band | Group | Number | \tilde{f} (mJy) | \tilde{L}_{FIR} ($10^{12} L_{\odot}$) | \tilde{L}_{IR} ($10^{12} L_{\odot}$) | $\langle f \rangle$ (mJy) | $\langle L_{\text{FIR}} \rangle$ ($10^{12} L_{\odot}$) | $\langle L_{\text{IR}} \rangle$ ($10^{12} L_{\odot}$) |
|------|---------------------------------|--------|----------------------|---|--|------------------------------|---|--|
| (1) | (2) | (3) | (4) | (5) | (6) | (7) | (8) | (9) |
| 850 | Whole sample | 50 | 1.93 ± 0.35 | 2.04 ± 0.37 | 2.87 ± 0.52 | 2.30 ± 0.16 | 2.43 ± 0.17 | 3.43 ± 0.24 |
| | 850 μm detections | 18 | 4.32 ± 0.54 | 4.56 ± 0.57 | 6.43 ± 0.80 | 5.00 ± 0.27 | 5.28 ± 0.28 | 7.44 ± 0.40 |
| | 850 μm nondetections | 32 | 0.80 ± 0.35 | 0.84 ± 0.37 | 1.19 ± 0.52 | 0.70 ± 0.20 | 0.74 ± 0.22 | 1.05 ± 0.31 |
| 450 | Whole sample | 50 | -0.22 ± 2.29 | 0.13 ± 1.06 | -0.18 ± 1.85 | 0.51 ± 1.82 | 0.29 ± 1.05 | 0.41 ± 1.47 |
| | 850 μm detections | 18 | 3.90 ± 3.52 | 2.25 ± 2.02 | 3.17 ± 2.85 | 4.39 ± 2.80 | 2.53 ± 1.61 | 3.56 ± 2.27 |
| | 850 μm nondetections | 32 | -2.64 ± 3.30 | 1.52 ± -1.14 | -2.14 ± 2.68 | -2.30 ± 2.39 | 1.32 ± -1.28 | -1.87 ± 1.94 |

NOTE—

(1) band; (2) the groups of the sample; (3) the number of quasars in each group; (4)(5)(6) the median value of flux density, FIR and IR luminosities; (7)(8)(9) the stacked average value with inverse-square variance weighting. Note here we exclude four radio-loud quasars which have $>3.5\sigma$ detections in FIRST survey.

SCUBA2 images at $450\ \mu\text{m}$ and $850\ \mu\text{m}$ with inverse-square variance weighting following the formula $F_s = \sum_{i=1}^n \frac{F_i/\sigma_i^2}{\sum_{i=1}^n 1/\sigma_i^2}$, $\sigma_s = \frac{1}{\sqrt{\sum_{i=1}^n 1/\sigma_i^2}}$, like some previous work of faint extra-galactic sources (e.g., [Violino et al. 2016](#)).

For detections we centered the image at the quasar sub-mm position; and for non-detections, we stacked at the quasar optical position. We list the results of both median and weighted average stacking results in [Table 3](#) and present the stacked maps in [Figure 3](#). The value of the central pixel was considered as the stacked flux density of each group. The error of the median was measured from the stacked median map. First we calculated the median value for every pixel for the samples to construct a median map. Then, we masked the center part of the map (i.e. the location of the quasar). Then the standard deviation value of the pixels on this source-masked map was considered as the error of the final median flux density.

We note that because our sample are all point sources, the offsets in 7/16 detections are caused by the large beam size. The peak pixel value is the total flux density of the point source. Thus, for detections we preferred to stack them as the sub-mm peak flux pixel. Here we excluded four radio-loud quasars (i.e., P055–00, P135+16, J1207+0630 and J1609+3041) which have $>3.5\sigma$ detections in FIRST survey, to avoid possible contamination from the radio jet in the FIR band.

We measured the average flux density for $850\ \mu\text{m}$ detections of $F_{s,850\ \mu\text{m}} = 5.00 \pm 0.27\ \text{mJy}$, $\text{SNR} = 18.5$ from the $850\ \mu\text{m}$ detected stacked map (in [Figure 3](#)), while for all $z \sim 6$ quasars $F_{s,850\ \mu\text{m}} = 2.30 \pm 0.16\ \text{mJy}$, $\text{SNR} = 14.4$. [Table 3](#) shows the median and weighted average parameters of $z \sim 6$ quasars in our survey. The FIR and infrared (IR) luminosities are calculated assuming a graybody at $z = 6$, as described in [Section 4](#). The average FIR properties in SHERRY are comparable with the average FIR luminosity of $2.0 \pm 0.3 \times 10^{12} L_\odot$ for $z \sim 6$ quasars using MAMBO-II in [Wang et al. \(2011b\)](#). The median FIR luminosity of our sample is also very similar to that of accretion-rate-limited $z \sim 6$ quasars of $1.8 \times 10^{12} L_\odot$ reported by [Venemans et al. \(2018\)](#) based on ALMA observaions.

4. SED FITTING AND FIR PROPERTIES

4.1. Spectral energy distributions of the SCUBA2 detections

The sub-mm/mm surveys reveal dust masses of a few $10^8 M_\odot$ in the host galaxies of the $z \sim 6$ quasars (e.g. [Omont et al. 2013](#); [Venemans et al. 2016, 2018](#)). The central AGN can heat the dust torus to a few hundred to $> 1,000\ \text{K}$ which dominates the near-IR and mid-IR

emission (e.g., [Leipski et al. 2013, 2014](#)). If there is active star formation in the quasar hosts, additional dust heating by the intense UV photons from OB stars, except AGN, will result in bright FIR continuum emission (i.e., dust temperatures of 30–80 K, e.g., [Beelen et al. 2006](#); [Wang et al. 2007](#)). For objects at $z \sim 5.6$ to 6.9, this could be traced by our SHERRY observations at $850\ \mu\text{m}$ at mJy sensitivity.

Our SHERRY observations detected twenty sources at $850\ \mu\text{m}$. We also collect the optical, near-IR, and mid-IR photometric data from SDSS, PS1, and *WISE*³, and plot the rest-frame UV-FIR SED of these objects in [Figure 4](#). In addition to the SCUBA2 measurements, we also included available sub-mm/mm data from IRAM 30 m/MAMBO and ALMA from the literature (e.g., [Venemans et al. 2018](#)). We first compared the SEDs to the templates of quasars at low redshift. We included two quasar templates in [Figure 4](#). One is the intrinsic AGN SED template, which is derived with the sample of $z < 0.18$ optically selected PG QSOs ([Symeonidis et al. 2016](#)). The other is the optical to mid-IR SED template derived from SDSS and *Spitzer* photometry of 259 optically luminous quasars ([Richards et al. 2006](#)). We interpolated the *W1* and *W2* data to the rest-frame $5100\ \text{\AA}$ assuming a power-law spectrum and scale the templates to this monochromatic luminosity. For the sources having *W3* band data, i.e. J0100+2802, J0148+0600, P135+16 and P215-16, we fit the power law using *W1*, *W2* and *W3* data and interpolate the *WISE* data to the rest-frame $5100\ \text{\AA}$.⁴

Our SHERRY observations, as well as the sub-mm/mm data from previous observations, suggest strong FIR continuum that exceed the AGN templates in nineteen objects (19/20 SCUBA2 detections), shown in [Figure 4](#). Combining these data, we fitted the dust emission with a modified blackbody model following the formula described in [De Breuck et al. \(2003\)](#):

$$\begin{aligned} L_{\text{IR}} &= 4\pi D_L^2 \int_0^\infty S_{\nu_r} d\nu_r \\ \Rightarrow L_{\text{IR}} &= 4\pi \Gamma[\beta + 4] \zeta[\beta + 4] D_L^2 x^{-(\beta+4)} (e^x - 1) S_{\nu_r} \nu_r \end{aligned} \quad (1)$$

³

SDSS catalog: [Alam et al. \(2015\)](#); [Abazajian et al. \(2009\)](#)

ALLWISE catalog: [Cutri & et al. \(2014\)](#)

The Pan-STARRS1 Surveys: [Chambers et al. \(2016\)](#)

⁴ Here we have compared the normalizing at $5100\ \text{\AA}$ and $1\ \mu\text{m}$ for these four sources with *W3* available. If we interpolate *WISE* data around $1\ \mu\text{m}$, the FIR excess are $[< 6.1, 0.9 \pm 1.8, 3.1 \pm 2.0, 6.4 \pm 8.8] \times 10^{13}\ \text{erg/s}$, respectively. The differences of FIR excess are within the error bar, which doesn't affect the major conclusions in the paper.

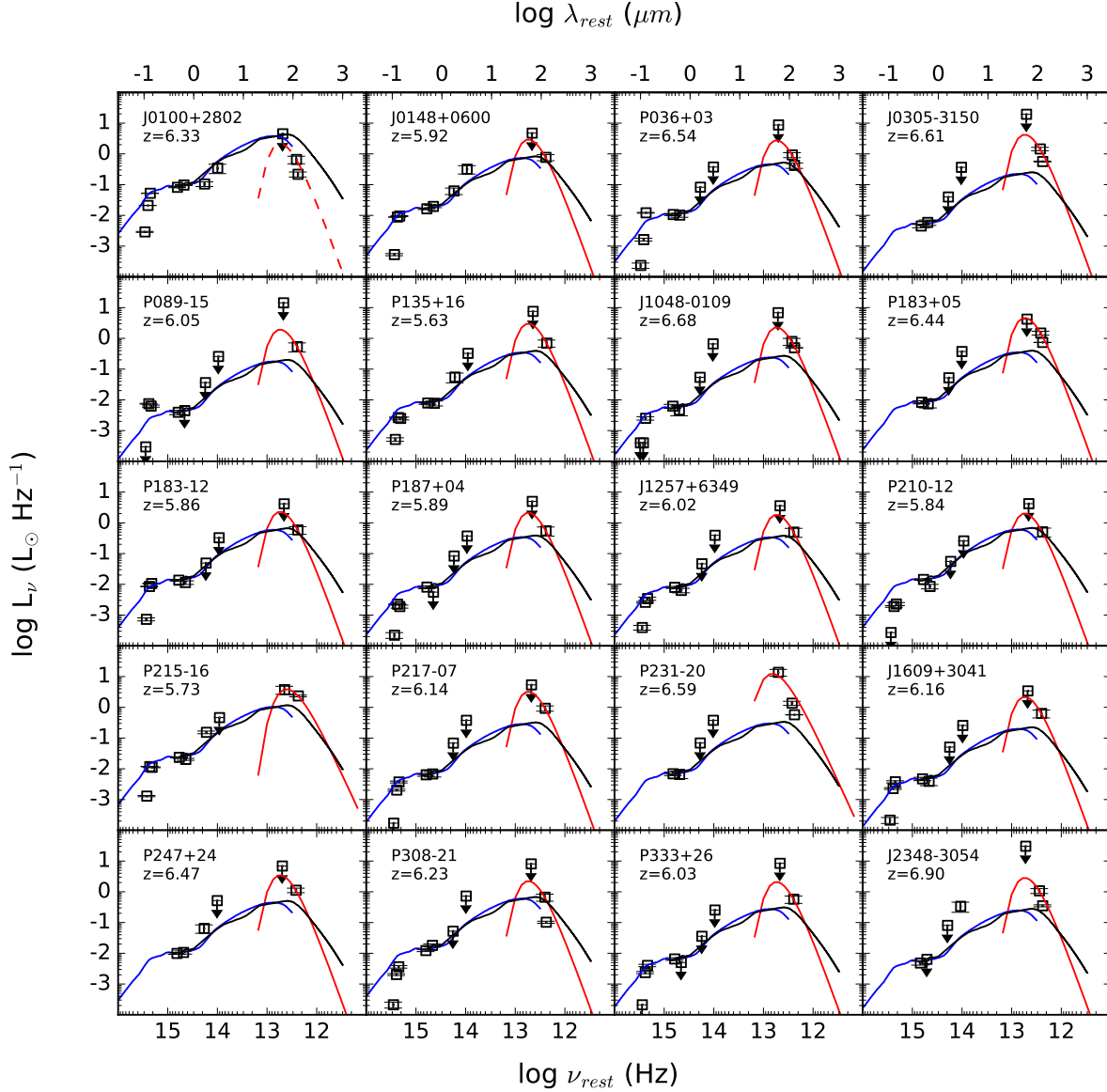


Figure 4. Individual UV-to-radio SEDs of the sub-millimeter detected $z \sim 6$ quasars. The points show the data from SDSS, PS1, *WISE*, PdBI, ALMA and SCUBA2 measurements, listed in Table 2. The arrows denote 3σ upper limits. Two AGN SED templates (Symeonidis et al. 2016; Richards et al. 2006) are plotted and scaled to 5100 \AA . The black line is the intrinsic AGN SED template derived from a sample of $z < 0.18$ unobscured and optical luminous PG QSOs (Symeonidis et al. 2016). The blue line is Type I quasar template derived from the *Spitzer* survey of SDSS quasar (Richards et al. 2006). The red lines are the graybody fitting to the FIR emission.

where S_{ν_r} is the rest-frame flux density, ν_r is rest frequency, D_L is the luminosity distance. T_{dust} is dust temperature in K, β is emissivity index. $x = h\nu/kT_{\text{dust}}$, Γ and ζ are the Gamma and Riemann ζ functions, respectively. $S_{\nu_r} \propto (1 - e^{-\tau_{\text{dust}}})\nu^3/(e^x - 1)$. Here we followed Beelen et al. (2006) and assumed that the dust is optically thin at far-infrared wavelengths, i.e., $\tau_{\text{dust}} \ll 1$, at $\lambda > 40 \mu\text{m}$. However, if an optical depth of $\tau_{\text{dust}} = 1$ is assumed for the dust emission detected at $850 \mu\text{m}$, the

derived FIR luminosities will be lower by a factor of ~ 3 with assumptions of $T_{\text{dust}} = 47 \text{ K}$ and $\beta = 1.6$.

Based on optically thin assumption, there are three free parameters to be fitted: L_{IR} , dust temperature T_{dust} , and emissivity index β . If the objects only have one data points (11/20) or two data points on one side of the gray body peak (7/20) at FIR band, which is not enough to fit the curve, we fixed $T_{\text{dust}} = 47 \text{ K}$, $\beta = 1.6$, which are the typical values found for FIR bright quasars

Table 4. The results and derived parameters of SHERRY detections

| Source name | L_{bol} ($10^{13} L_{\odot}$) | \dot{M} ($M_{\odot} \text{ yr}^{-1}$) | L_{FIR} ($10^{12} L_{\odot}$) | $L_{\text{IR}} (L_{\text{IR,SF}})$ ($10^{12} L_{\odot}$) | M_{dust} ($10^8 M_{\odot}$) | SFR ($10^2 M_{\odot} \text{ yr}^{-1}$) |
|------------------|---|--|---|---|---|---|
| (1) | (2) | (3) | (4) | (5) | (6) | (7) |
| J0100+2802 | 47.2 | 318.5 | 4.3 ± 1.2 | 6.1 ± 1.7 (<6.1) | 2.4 ± 0.7 | <6.1 |
| J0148+0600 | 8.2 | 55.4 | 5.6 ± 1.3 | 7.9 ± 1.8 (2.1 ± 1.8) | 3.2 ± 0.7 | 2.1 ± 1.8 |
| J036+03 | 8.2 | 55.4 | 5.7 ± 1.1 | 8.1 ± 1.5 (4.1 ± 1.5) | 3.3 ± 0.6 | 4.1 ± 1.5 |
| J0305–3150 | 2.3 | 15.2 | 8.9 ± 1.1 | 12.5 ± 1.6 (10.2 ± 1.6) | 5.0 ± 0.6 | 10.2 ± 1.6 |
| J089–15 | 5.7 | 38.3 | 3.8 ± 1.2 | 5.3 ± 1.7 (3.6 ± 1.7) | 2.1 ± 0.7 | 3.6 ± 1.7 |
| J135+16 | 2.5 | 16.7 | 5.5 ± 1.4 | 7.8 ± 2.0 (4.6 ± 2.0) | 3.1 ± 0.8 | 4.6 ± 2.0 |
| J1048–0109 | 2.5 | 16.7 | 4.8 ± 1.2 | 6.7 ± 1.7 (4.5 ± 1.7) | 2.7 ± 0.7 | 4.5 ± 1.7 |
| J183+05 | 6.2 | 42.0 | 9.5 ± 1.4 | 13.4 ± 1.9 (9.9 ± 1.9) | 5.4 ± 0.8 | 9.9 ± 1.9 |
| J183–12 | 8.2 | 55.4 | 4.3 ± 1.2 | 6.1 ± 1.7 (1.5 ± 1.7) | 2.5 ± 0.7 | 1.5 ± 1.7 |
| PSOJ187+04 | 1.4 | 9.6 | 4.0 ± 1.3 | 5.7 ± 1.9 (2.7 ± 1.9) | 2.3 ± 0.8 | 2.7 ± 1.9 |
| SDSSJ1257+6349 | 4.3 | 29.0 | 3.5 ± 1.1 | 5.0 ± 1.6 (2.1 ± 1.6) | 2.0 ± 0.6 | 2.1 ± 1.6 |
| PSOJ210–12 | 1.9 | 12.7 | 3.8 ± 1.2 | 5.3 ± 1.7 (0.9 ± 1.7) | 2.1 ± 0.7 | 0.9 ± 1.7 |
| PSOJ215–16 | 10.8 | 73.0 | 14.1 ± 6.6 | 18.8 ± 8.8 (10.6 ± 8.8) | 10.2 ± 4.8 | 10.6 ± 8.8 |
| PSOJ217–07 | 3.6 | 24.2 | 6.4 ± 1.2 | 9.0 ± 1.7 (6.2 ± 1.7) | 3.6 ± 0.7 | 6.2 ± 1.7 |
| P231–20 | 7.5 | 50.5 | ~ 30.0 | ~ 71.8 (<68.2) | ~ 7.9 | <68.2 |
| SDSSJ1609+3041 | 4.3 | 29.0 | 4.3 ± 1.2 | 6.1 ± 1.7 (4.2 ± 1.7) | 2.4 ± 0.7 | 4.2 ± 1.7 |
| P247+24 | 4.3 | 29.0 | 7.1 ± 1.2 | 10.0 ± 1.7 (5.9 ± 1.7) | 4.0 ± 0.7 | 5.9 ± 1.7 |
| PSOJ308–21 | 3.6 | 24.2 | ~ 3.9 | ~ 5.5 (~ 1.4) | ~ 2.2 | ~ 1.4 |
| PSOJ333+26 | 3.6 | 24.2 | 4.0 ± 1.1 | 5.7 ± 1.5 (3.2 ± 1.5) | 2.3 ± 0.6 | 3.2 ± 1.5 |
| VIKINGJ2348–3054 | 2.1 | 13.9 | 6.2 ± 1.1 | 8.7 ± 1.6 (6.3 ± 1.6) | 3.5 ± 0.6 | 6.3 ± 1.6 |

Notes.

- The dust mass and SFR of these quasars are calculated from IR luminosity after removing the AGN contribution.
- We adopt an emissivity index of $\beta = 1.6$, $T_{\text{dust}} = 47 \text{ K}$ here for all the calculations (Beelen et al. 2006).
- The AGN bolometric luminosities are estimated by the UV luminosities (1450 \AA) with $L_{\text{bol}} = 4.2\nu L_{\nu,1450}$ (Runnoe et al. 2012), and then we convert L_{bol} to black hole accretion rate as $L_{\text{bol}} = \eta \dot{M} c^2$ assuming the efficiency $\eta = 0.1$.
- ALMA observations resolved two sources, P23120 and PSO J30821, having a companion in a SCUBA2 beam (Decarli et al. 2017); here we derived their infrared luminosities and other relevant parameters using the ALMA continuum flux ratio between the quasar and its companion.

at lower redshift $z \sim 2-4$ (Beelen et al. 2006). For the objects with two or more data points (2/20, P215–16 and P231–20), we fixed $\beta = 1.6$. For P215–16, the best parameters of $[L_{\text{IR}}, T_{\text{dust}}]$ are $[1.9 \pm 0.9 \times 10^{13} L_{\odot}, 42.4 \text{ K}]$ when we fix $\beta = 1.6$. And for P231–20, fixed $\beta = 1.6$, the best parameters are $[L_{\text{IR}}, T_{\text{dust}}] = [7.2 \times 10^{13} L_{\odot}, 70.0 \text{ K}]$. This dust temperature is higher than that of the high- z quasar (Leipski et al. 2014). It suggests that the AGN contribution is significant to the dust heating. The companion sources also introduce large uncertainties in IR luminosity calculation. Thus we consider the derived SFR from this IR luminosity as an upperlimit (Table 4). Two sources in our survey (P231–20 and P308–21) are reported having millimeter continuum companions ⁵ in

a SCUBA2 beam (Decarli et al. 2017). Here we derived their infrared luminosities and other relevant parameters using the ALMA continuum flux ratio between the quasar and its companion. The derived parameters for all detected sources are listed in Table 4.

The CMB temperature is $\sim 19.1 \text{ K}$ at $z = 6.0$. We checked the CMB effect following the description in Venemans et al. (2016). For galaxies with T_{dust} in the range of 42 to 70 K, the increase of dust temperature heated by the CMB is only 0.21–0.01%, which is negligible (Equation 4 of Venemans et al. 2016, see also da Cunha et al. 2013). The missing fraction of the dust continuum due to the CMB effect can be estimated as $S_{\nu}^{\text{obs}}/S_{\nu}^{\text{intrinsic}} = 1 - B_{\nu}[T_{\text{CMB}}(z)]/B_{\nu}[T_{\text{dust}}]$. With an assumption of the dust temperature of $T_{\text{dust}} = 47 \text{ K}$ and a redshift in a range of $z = [5.6-6.9]$, we are only missing 2.0–3.2% and 0.1–0.2% of the intrinsic flux density at observed wavelengths of $850 \mu\text{m}$ and $450 \mu\text{m}$, respectively. If we assume a dust temperature of $T_{\text{dust}} = 30 \text{ K}$ instead of 47 K, the missing fraction is 8.0–15.3% at $850 \mu\text{m}$ and 0.9–2.9% at $450 \mu\text{m}$ at observed wavelength. Considering that the measurement uncertainties are much larger than the CMB corrections in the tem-

⁵ Decarli et al. (2017) presents the J-band magnitude of companion is much fainter than the quasar. For P231-20, $m_{\text{AB,QSO}} = 19.66 \pm 0.05 \text{ mag}$, $m_{\text{AB,comp}} > 21.29 \text{ mag}$; for P308-21, $m_{\text{AB,QSO}} = 20.17 \pm 0.11 \text{ mag}$, $m_{\text{AB,comp}} > 21.89 \text{ mag}$. They also reported the separations between the quasars (P231-20, P308-21) and their companion sources are $1.6''$ and $2.4''$. For P231-20, the VLT long-slit width is $1.3''$ and Magellan slit width is $0.6''$. For P308-21, the VLT slit width is $1.3''$. Thus, the companion is unlikely to affect the measurement of NIR spectra and rest-frame UV line analysis.

perature range we adopt in this paper, we neglect the CMB effects and directly use the observed flux densities in the analysis and discussions below.

4.2. FIR luminosities and star formation rates

Integrating the FIR emission between 42.5 and 122.5 μm in the rest-frame allows us to determine its FIR luminosity (Helou et al. 1985; widely used in the papers on high- z quasar, e.g. Wang et al. 2007; Omont et al. 2013; Venemans et al. 2018).

$$\begin{aligned} L_{\text{FIR}} &= 4\pi M_{\text{dust}} \int \kappa(\nu) B_{\nu}(T_{\text{dust}}) d\nu \\ S_{\nu} &= \alpha \nu^{3+\beta} \frac{1}{e^{h\nu/KT_{\text{dust}}} - 1} \\ \Rightarrow M_{\text{dust}} &= \frac{c^2 D_L^2}{2h} \frac{\alpha \nu_0^{\beta}}{\kappa_0} \end{aligned} \quad (2)$$

where B_{ν} is the Planck function, h is the Planck constant, and $\kappa(\nu) = \kappa_0(\nu/\nu_0)^{\beta}$ is the dust absorption coefficient. We adopt $\kappa_0 = 18.75 \text{ cm}^2 \text{ g}^{-1}$ at 125 μm (Hildebrand 1983). The derived dust masses are in the range of 2.0 – 10.2 $\times 10^8 M_{\odot}$.

Wang et al. (2008a) found the FIR luminosities around $\sim 10^{13} L_{\odot}$ with the warm dust temperatures of 39–52 K in four $z > 5$ SDSS quasars using SHARC-II at 350 μm . Later, Leipski et al. (2013) also reported the FIR emission of $\sim 10^{13} L_{\odot}$ in 69 QSOs at $z > 5$ with the cold component temperature of ~ 50 K. The FIR luminosities in our SCUBA2 survey are around 0.4 to 3.0 $\times 10^{13} L_{\odot}$, which is close to the previous results at $z > 5$ (e.g., Wang et al. 2008a, 2013; Willott et al. 2013; Venemans et al. 2016).

In addition to the dust emission powered by the central AGN, a similar excess of FIR emission heated by host galaxy star formation was widely reported with the samples of FIR-mm detected quasars from low- z to high- z . At low redshift, Shangguan et al. (2018) studied 87 $z < 0.5$ PG quasars and found the minimum radiation field intensity of the galaxy increases with increasing AGN luminosity (Figure 6(a) in Shangguan et al. 2018), which imply the quasars can heat dust on galactic scales. Symeonidis et al. (2016) used 3 luminous QSO samples from the literatures to compare the FIR excess, i.e. Type I radio-quiet QSOs with robust submm/mm detections at $1.7 < z < 2.9$ from Lutz et al. (2008), 24 μm -selected broad-line QSOs at $1.7 < z < 3.6$ from Dai et al. (2012) and X-ray absorbed and submm-luminous Type I QSOs at $1.7 < z < 2.8$ from Khan-Ali et al. (2015). The results shows the FIR excess is dominant if the intrinsic AGN power at 5100 \AA is more than a factor of 2 lower than the galaxy's 60 μm luminosity and more than factor of 4 lower than the total IR emission (8-1000 μm) of the galaxy. At high redshift, Wang

et al. (2008a) also reported nine 250GHz detected $z \sim 6$ quasars (9/10) having significant IR excess components, tracing the dust heated by the star formation activities in the host galaxies. To constrain SFRs from the FIR excess, the contribution from AGN should be estimated and removed.

Here we found nineteen SCUBA2 detections (19/20) have FIR excess, which is also seen before with other quasar samples. We calculated SFR derived by IR excess component (8–1000 μm), which is corrected by removing the contribution of the AGN using the intrinsic AGN template (Symeonidis et al. 2016). We converted IR luminosity into star-formation rate using the formula $\text{SFR}(M_{\odot} \text{ yr}^{-1}) = 1.0 \times 10^{-10} L_{\text{IR}}(L_{\odot})$ assuming a Chabrier IMF, like some previous work (e.g., Magnelli et al. 2012). The estimated $\text{SFR} = 90\text{--}1060 M_{\odot} \text{ yr}^{-1}$ with IR excess luminosity, which is consistent with previous works for $z \sim 6$ quasars (e.g., Wang et al. 2008b; Venemans et al. 2018; Wang et al. 2013). For SDSS J0100+2802, due to a possibly different β or T_{dust} , our SED fitting is below the AGN template. Its SFR calculated by total IR luminosity is considered as the upper limit in the quasar host galaxy. More mid-IR observations for this source are required.

4.3. The redshift evolution of AGN properties

4.3.1. UV to FIR SED and comparisons to lower redshift

We adopt the median stacking approach and calculate the median SED for our SHERRY $z \sim 6$ quasars. For the SCUBA2 median maps described in Section 3.2, if the median value at the quasar position was larger than three times that of the background, we considered the median quasar signal to be significant; otherwise the signal was considered an upper limit. The median parameters are listed in Table 3. For the other photometry data, to estimate the median and its variation, we considered the error bar of each individual data point and used a bootstrapping approach. The data were selected as many times as the size of a given sample, allowing for replacements, to create a new sample to calculate a median. Note the selected data is a random value produced by a normal distribution generated by the measured data, considering error bars for every point. This process was repeated for 1,000 times, then we fitted a normal distribution to these 1,000 individual median value. The centroid of the distribution was the final median flux of this sample, and the standard deviation of this distribution was the uncertainty on the median flux. If the number of upper limits or non-detections was larger than 30% of the total sample number (e.g., W3 and W4) we considered the median value as an upperlimit. The resulting median SED data were stacked

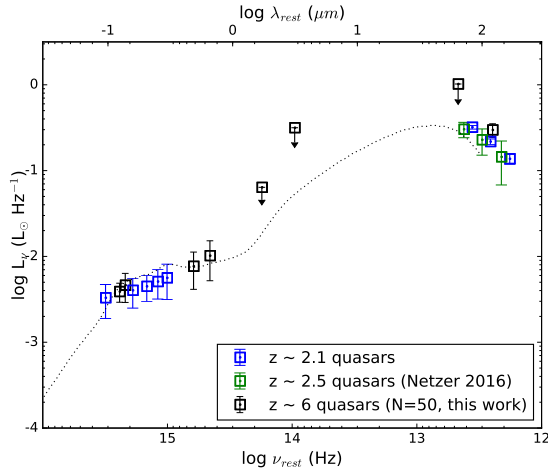


Figure 5. The median SED of a sample of quasars at $z \sim 6$ (black) and that of $z \sim 2.1$ quasars (blue). The dotted line is the SDSS quasar template (Richards et al. 2006), which is scaled to $W1$ band of the median SED of $z \sim 6$. Here we also exclude the radio-loud quasars.

at the observed frame and then shifted into rest-frame using the typical redshift ($z = 6$).

To study the redshift evolution of optical-to-IR SED, we also selected a sample of Type-I radio quiet quasars at $z \sim 2.1$ from the SDSS DR14 catalog⁶ that have similar 1450 Å luminosities with our $z \sim 6$ sample. We collected available SDSS data and preserve $z \sim 2.1$ quasars within the coverage of *Herschel* SPIRE image data from all sky database (Griffin et al. 2010), *Herschel* High Level Images in IPAC (Poglitsch et al. 2010). Additionally, targets were excluded if they were located in the edges of the image or crowded region like galaxy cluster as well were affected by gravitational lensing or strong galaxy interaction. Next, these preserved targets were sampling again to make the luminosity distribution comparable to the distribution of $z \sim 6$ sample to avoid luminosity bias. Moreover, we combined all the images of resampled targets at each SPIRE band to the medium stacked image. Finally a stacking SED was estimated from flux density measured by SPIRE point source photometry pipeline. In Figure 5, we compared the median SED of a quasar sample at $z \sim 6$ with the low- z comparison sample at $z \sim 2.1$. We scaled a Type I quasar template, derived from the *Spitzer* survey of SDSS quasars (Richards et al. 2006), to $W1$ band luminosity of the median SED of the $z \sim 6$ quasar sample. The continuum emission measured by SPIRE at $350\mu\text{m}$ for the $z \sim 2.1$ sample is close to that measured by SCUBA2 at $850\mu\text{m}$ for the $z \sim 6$ sample in the rest

frame. We also compared with the median stack *Herschel*/SPIRE data of 100 luminous, optically selected active galactic nuclei (AGNs) at $z = 2 - 3.5$ (Netzer et al. 2016), shown as the green squares. The median SED at $z \sim 6$ is similar with these low- z comparison sample at $z \sim 2 - 3.5$, which implies these AGNs have similar dust emission properties and broad-band continuum emission.

4.3.2. FIR luminous quasars are fewer at $z \sim 6$

Figure 6(a) shows the relation between FIR luminosity and redshift for the $z \sim 6$ sample and all the comparison samples. The $z \sim 6$ quasars are from Wang et al. (2011a); Omont et al. (2013); Decarli et al. (2018); Venemans et al. (2018) and our work; the lower redshift AGNs are MAMBO-250 GHz or SCUBA-350 GHz observations of optically bright quasars at $z \sim 2-5$ (Carilli et al. 2001; Omont et al. 2001; Omont et al. 2003; Priddey et al. 2003). The FIR luminosities for detected or non-detected samples are shown by the filled or open symbols. The L_{FIR} values of all these high redshift quasars lie between $10^{11.4}$ and $10^{13.6} L_{\odot}$; for $z \sim 6$ quasars, FIR values are in the range of $10^{11.4} - 10^{13.3} L_{\odot}$ with the mean value of $10^{12.6} L_{\odot}$. For the bright tails of the FIR luminosity, it can be seen that FIR luminous quasars (e.g. $L_{\text{FIR}} > 10^{13.0} L_{\odot}$) are less common at $z \sim 6$ compared with the lower redshift ones. For the $z \sim 2$ sample, the average FIR luminosity for the objects that are detected at sub-mm/mm is $10^{13.2} L_{\odot}$. At $z \sim 6$, we didn't see a significant population with FIR luminosity at this level. This suggests that $z \sim 6$ quasar hosts are less evolved compared to the most luminous quasars at $z \sim 2$.

4.3.3. FIR luminosities and AGN luminosities

The relationship between FIR luminosity and AGN luminosity is $L_{\text{FIR}} = L_{\text{AGN}}^{0.7-0.8}$ for typical optically-selected Palomar-Green quasars and $L_{\text{FIR}} = L_{\text{AGN}}^{0.4}$ for the local IR-luminous quasars hosted by starburst ULIRGs (Hao et al. 2005). The high redshift quasars with bright submm/mm detections also follow the correlation of the local IR-luminous quasars, which suggest extreme starburst in their host (Wang et al. 2007, 2011a; Lutz et al. 2010). Dong & Wu (2016) studied $z < 4$ SDSS quasars in the *Herschel* Stripe 82 survey and presented the FIR-to-AGN luminosity relation as $L_{\text{FIR}} \propto L_{\text{bol}}^{0.46 \pm 0.03}$. Here we also investigate the FIR-to-AGN luminosity correlation in our $z \sim 6$ sample, and compare with the sample of MAMBO-250 GHz or SCUBA-350 GHz observed quasars at $z \sim 2-5$ (Omont et al. 2001; Omont et al. 2003; Carilli et al. 2001; Priddey et al. 2008). In this work, the AGN bolometric luminosities is estimated by the UV luminosities (1450 Å) with $L_{\text{bol}} =$

⁶ SDSS DR14 catalog: <https://www.sdss.org/dr14/>

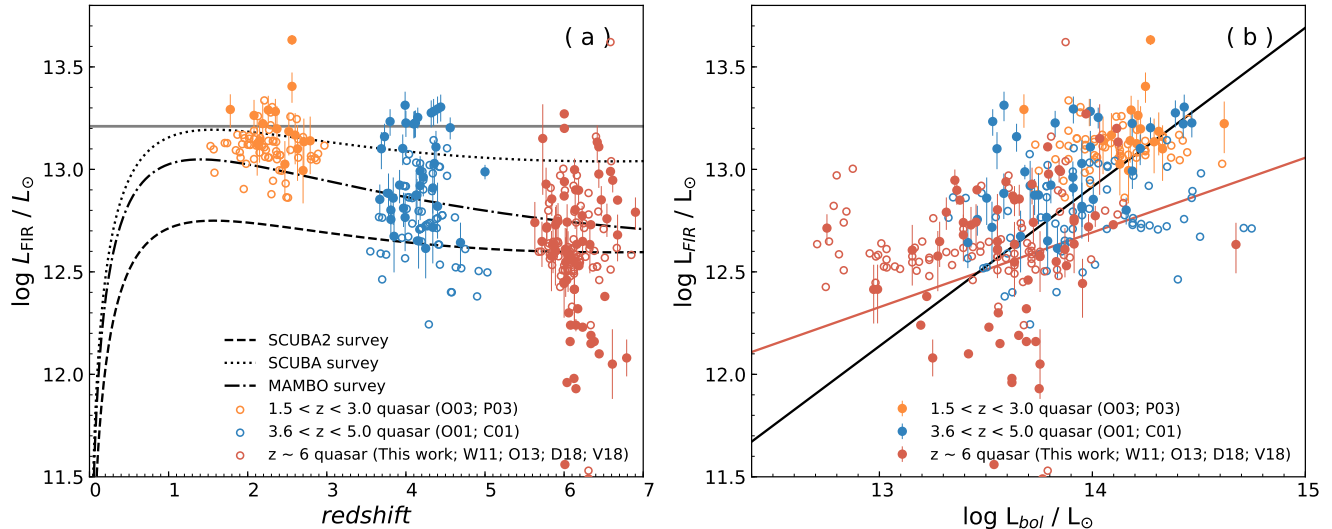


Figure 6. (a) The relation between FIR luminosity and redshift. The red points include almost all published mm-observed $z \sim 6$ quasars, from Wang et al. (2011a); Omont et al. (2013); Decarli et al. (2018); Venemans et al. (2018) (W11; O13; D18; V18) and this work. The $1.5 < z < 3.0$ group (orange) is combined with the samples from Omont et al. (2003) and Priddey et al. (2003) (O03; P03), while the $3.6 < z < 5.0$ group (blue) is from Omont et al. (2001) and Carilli et al. (2001) (O01; C01). The filled symbols represent detections with 1σ errors, and the open symbols indicate 3σ upper limits of the non-detections. The dashed and dotted lines represent the typical 3σ detection limits of SCUBA2 and SCUBA at 350 GHz respectively, $S_{350\text{ GHz}} = 3.6$ mJy (SCUBA2; in this work) and $S_{350\text{ GHz}} = 10$ mJy (Priddey et al. 2003). The dotted dashed line represents the typical 3σ detection limits of MAMBO survey at 250 GHz with $S_{250\text{ GHz}} = 2.4$ mJy (Wang et al. 2011a). The solid red points below the SCUBA2 detection limit are from recent ALMA observations (e.g. Decarli et al. 2018; Venemans et al. 2018), with a much better sensitivity. (b) The relationship of FIR luminosity and bolometric luminosity for different redshift sample. The black line represents the linear fitting for the detections for all redshift sample; the red line indicates the fitting for $z \sim 6$ sample.

$4.2\nu L_{\nu,1450}$ (Runnoe et al. 2012). We also recalculate the bolometric luminosities in other papers (e.g., Wang et al. 2011a; Omont et al. 2013) with the same conversion factor from Runnoe et al. (2012).

Figure 6(b) shows the far-infrared and AGN bolometric luminosity correlations in different redshift groups. For the non-detection of all samples, we also calculate the 3σ upper limits, marked as open symbols. In Figure 6(b), we fit the relation of L_{FIR} and L_{bol} for $z \sim 6$ and all high redshift detected quasars using least squares polynomial as follow:

$z \sim 6$ detected quasars :

$$\log L_{\text{FIR}} = (0.36 \pm 0.14) \log L_{\text{bol}} + (7.59 \pm 1.84) \quad (3)$$

All detections :

$$\log L_{\text{FIR}} = (0.78 \pm 0.04) \log L_{\text{bol}} + (2.04 \pm 0.52)$$

The Pearson correlation coefficient r -value for $z \sim 6$ detected sample is 0.317 with p -value = 0.009, which implies a correlation between them. Wang et al. (2007) reported there is no correlation between optical and FIR luminosities for the $z \sim 6$ sample, which is mainly due to the narrow luminosity range and small sample size. The r -value is 0.81 when including all the sample, which argues for a correlation. From Figure 6(b), the FIR and

bolometric luminosities of the optically selected quasars from local to high- z show a correlation with large scatters. At $\log L_{\text{bol}} / L_{\odot} \sim 13.7$, the red dots span ~ 1.2 dex. The correlation suggests connection between the two parameters; the FIR emitting dust in the nuclear region could be heated by both AGN and star formation, the star formation and SMBH accretion are fueled by the same gas reservoir and correlated to the mass of the host galaxies (e.g., Xu et al. 2015). And the scatters imply that at a given bolometric luminosity, quasars could show a range of FIR luminosities heated by different level of star forming activities in the host as was discussed in a range of papers (e.g., Schulze et al. 2019; Netzer et al. 2016). For example, the scatter in previous MAMBO survey for $z \sim 6$ quasars in Wang et al. (2008b) is about 1 dex at $\log L_{\text{bol}} / L_{\odot} \sim 13.9$. The luminosities and upper limits of the $z \sim 6$ quasars extend the L_{FIR} -to- L_{bol} trends of the two local samples and mix them at the high luminosity end (Wang et al. 2008b). Such scatters were also reported with submm/mm observations of quasars at lower redshift, e.g. Dong & Wu (2016) found that the L_{FIR} -to- L_{bol} relation of $z < 4$ SDSS quasars in the *Herschel* Stripe 82 survey have a large scatter of ~ 1 dex. The objects that are luminous in the optical with extreme starburst in the quasar host galaxies mark the

upper boundary of this FIR-quasar bolometric luminosity correlation. This is shown with the FIR luminous quasars in the $z \sim 2$ to 6 samples in Figure 6(b).

5. WEAK LINE FEATURES

Diamond-Stanic et al. (2009) studied > 5000 quasars at $z > 3$ selected down to a magnitude limit of $i = 20.2$ mag in the SDSS DR5 quasar catalog. They defined 74 weak-line quasars at $z > 3$ as the ones that have a rest-frame equivalent width (EW) of the $\text{Ly}\alpha + \text{Nv}$ line (determined between $\lambda_{\text{rest}} = 1160 \text{ \AA}$ and $\lambda_{\text{rest}} = 1290 \text{ \AA}$) lower than 15.4 \AA , while the mean value is 62 \AA for the normal SDSS quasars. Bañados et al. (2016) reported that objects with such weak line took about 13.7% in the sample of 124 quasars at $5.6 < z < 6.7$ discovered from PS1. People also reported connections between the weak line feature in quasar UV spectra and sub-mm dust continuum detections (Omont et al. 1996; Bertoldi et al. 2003a; Wang et al. 2008b). Wang et al. (2008b) presented mm observations using IRAM/MAMBOII for eighteen $z > 5.7$ quasars, and showed mm detections tended to have weaker $\text{Ly}\alpha$ emission than the non-detected sources (see Figure 5 in Wang et al. 2008b). Following these ideas, the SHERRY survey extends the samples to study the link between the FIR properties and the UV emission line at $z \sim 6$.

Appendix A shows the spectra of all of the 54 quasars in our sample. These spectra include 52 published spectra provided by the authors of their discovery papers (see Table 5) and 2 unpublished spectra (Bañados et al. in prep., S.J. Warren et al. in prep.). The instruments and spectral resolution of these NIR spectra are summarized in Appendix B. For each spectrum, we fit a power law of $f_\lambda = C \times \lambda^\beta$ to the continuum and measure $\text{EW}(\text{Ly}\alpha + \text{Nv})$ following the procedure in Diamond-Stanic et al. (2009). The derived $\text{EW}(\text{Ly}\alpha + \text{Nv})$ is listed in Table 5. Figure 7 shows the log-normalized distribution of EW in our sample. The fraction of weak line quasars is 11.1% (6/54) according to the definition in Diamond-Stanic et al. (2009). The best fit is $\mu(\log \text{EW} / \text{ \AA}) = 1.658$ and $\sigma(\log \text{EW} / \text{ \AA}) = 0.385$ (black line), which is consistent with previous results of $\mu(\log \text{EW} / \text{ \AA}) = 1.542$ and $\sigma(\log \text{EW} / \text{ \AA}) = 0.391$ from PS1 sample at $z > 5.6$ in Bañados et al. (2016) (black dashed line). Diamond-Stanic et al. (2009) found that the best fit is $\mu(\log \text{EW} / \text{ \AA}) = 1.803$ and $\sigma(\log \text{EW} / \text{ \AA}) = 0.205$ for $3 < z < 5$ SDSS quasars (blue dashed line). $\text{EW}(\text{Ly}\alpha + \text{Nv})$ distribution at high redshift has a lower peak and a larger dispersion, which is also suggested by Bañados et al. (2016). Bañados et al. (2016) also pointed out that this could be the stronger IGM absorption at $z > 5.6$, or an indication of the

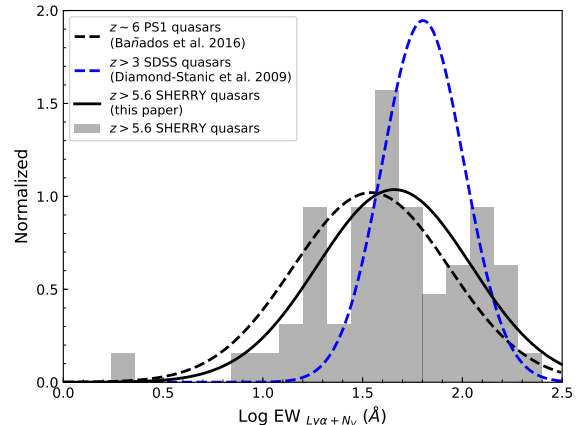


Figure 7. Log-normalized distribution of the $\text{EW}(\text{Ly}\alpha + \text{Nv})$. The black line represents the best fit to $z > 5.6$ quasars in our survey, with $\mu(\log \text{EW} / \text{ \AA}) = 1.658$ and $\sigma(\log \text{EW} / \text{ \AA}) = 0.385$, which is consistent with previous results of $\mu(\log \text{EW} / \text{ \AA}) = 1.542$ and $\sigma(\log \text{EW} / \text{ \AA}) = 0.391$ from PS1 sample at $z > 5.6$ in Bañados et al. (2016), shown in black dashed line. The blue dashed line shows the best fit for $3 < z < 5$ SDSS quasars (Diamond-Stanic et al. 2009).

EW distribution evolution with redshift. We note that the luminosity ranges are slightly different between the $z \sim 6$ and $3 < z < 5$ samples. The quasars in Diamond-Stanic et al. (2009) are selected down to a magnitude limit of $i = 20.2$ mag at $z > 3$, corresponding to $L_{\text{bol}} > 5.3 \times 10^{13}$ erg/s. This is a little higher than the L_{bol} limit at $z \sim 6$ in Bañados’s sample and our sample of $> 1.0 \times 10^{13}$ erg/s. Possible redshift evolution of EW could be further tested with low- z quasar samples in a luminosity range comparable to that of the $z \sim 6$ sample, though it is beyond the goal of this paper.

5.1. The connection between FIR and Weak Line Features

Some scenarios have been proposed to explain the nature of weak-line quasars in many previous works, i.e., Bañados et al. (2014); Wang et al. (2008b); Luo et al. (2015); Shemmer & Lieber (2015); Bañados et al. (2016). For example, Hryniewicz et al. (2010) suggested that WLQs may represent an early stage of quasar evolution with different physical conditions of the broad emission line region. The IRAM/MAMBOII survey implied a weak trend between FIR luminosity and optical weak line feature (Wang et al. 2008b). But is there a physical connection between them? Here we revisit this issue by including the new $z \sim 6$ quasar sample from our SCUBA-2 observations.

Table 5. EW of Ly α + N v

| Source | Type ^a | EW (Ly α + N v) (\AA) | EW (Ly α + N v) ^c (\AA) | Ref. | Source | Type ^a | EW (Ly α + N v) (\AA) | EW (Ly α + N v) ^c (\AA) | Ref. |
|-------------------|--------------------|--|---|--------|-------------------|-------------------|--|---|------|
| (1) | (2) | (3) | (4) | (5) | (1) | (2) | (3) | (4) | (5) |
| J0008–0626 | | 151.6 | 78 | (1) | P183+05 | DLA | 45.5 | | (11) |
| P002+32 | | 110.6 | | (2) | P183–12 | WLQ | 20.8 | 11.8 | (14) |
| P007+04 | WLQ | 21.2 | | (2) | P184+01 | | 51.3 | | (2) |
| J0100+2802 | WLQ | 14.9 | 10 | (3) | P187+04 | | 20.1 | | (2) |
| J0148+0600 | LoBAL ^a | 96.0 | >87 | (1) | J1243+2529 | | 69.5 | | (8) |
| P036+03 | | 27.9 | | (2) | J1257+6349 | | 36.3 | 18 | (1) |
| J0305–3150 | | 17.0 | | (4) | P210+27 | | 48.6 | | (2) |
| P055–00 | | 28.1 | | (2) | P210+40 | | 102.8 | | (2) |
| P056–16 | | 153.8 | | (2) | J1403+0902 | | 8.6 | 8 | (1) |
| P060+24 | | 53.7 | | (2) | P210–12 | WLQ | 20.4 | 10.7 | (14) |
| P065–19 | | 247.0 | | (2) | P215–16 | BAL ^a | 85.9 | 109.5 \pm 83.1 | (15) |
| P089–15 | | 123.4 | | (2) | P215+26 | BAL ^a | 91.8 | | (13) |
| J0810+5105 | | 53.8 | | (8) | P217–16 | | 20.5 | | (2) |
| J0828+2633 | | 38.4 | | (5) | P217–07 | | 19.5 | | (2) |
| J0835+3217 | | 80.1 | | (8) | P231–20 | | 2.0 | | (11) |
| J0839+0015 | | 46.7 | | (6) | P239–07 | | 38.6 | | (2) |
| J0842+1218 | | 80.7 | 44 | (1) | J1609+3041 | | 30.9 | | (8) |
| J0850+3246 | | 13.8 | 10 | (1) | P247+24 | | 121.6 | | (11) |
| P135+16 | WLQ | 23.7 | | (2) | P261+19 | | 53.8 | | (11) |
| P159–02 | | 111.8 | | (2) | P308–21 | | 46.5 | | (2) |
| J1048–0109 | | 11.3 | | (7) | J2100–1715 | | 46.1 | | (12) |
| P167–13 | | 36.4 | | (2) | (2) P323+12 | | 125.0 | | (11) |
| J1143+3808 | | 30.8 | | (2)(8) | P333+26 | | 43.8 | | (2) |
| J1148+0702 | | 186.4 | | (8) | P338+29 | | 126.1 | | (2) |
| J1152+0055 | BAL ^a | 54.1 | | (9) | P340–18 | | 159.3 | | (2) |
| J1205–0000 | BAL ^a | –3.1 | | (10) | J2348–3054 | BAL ^a | 38.1 | | (6) |
| J1207+0630 | | 44.6 | 31 | (1) | P359–06 | | 34.9 | | (2) |

Notes.

- a. BALs are excluded in EW statistic because of their large uncertainty.
- b. P183+05 is a metal-poor proximate DLA with absorbing Ly α (Bañados et al. 2019). We also excluded it from statistic.
- c. The EW (Ly α + N v) is from the References.
- d. SCUBA-2 detections are marked in boldface.

References. (1) Jiang et al. (2015); (2) Bañados et al. (2016); (3) Wu et al. (2015); (4) Venemans et al. (2013); (5) S. J. Warren et al. (in prep.); (6) Venemans et al. (2015a); (7) Wang et al. (2017); (8) Jiang et al. (2016); (9) Izumi et al. (2018); (10) Matsuoka et al. (2016); (11) Mazzucchelli et al. (2017); (12) Willott et al. (2010); (13) E. Bañados et al. (in prep.); (14) Bañados et al. (2014); (15) Morganson et al. (2012).

We firstly excluded BALs, see Table 5. Some BALs shows a strong but non-Gaussian emission line, e.g. PSO J215–16, which may due to the outflow blowing the dust along the line of sight and the quasar is naked. Therefore, the equivalent width of emission line for BALs has a large bias to the statistic study. We also excluded P183+05 that has a DLA in front of it absorbing Ly α (Bañados et al. 2019).

5.1.1. *FIR bright quasars tend to have lower EW (Ly α + N v)*

Figure 8 shows the histograms of EW (Ly α + N v) for all SCUBA2 detections and non-detections. The best fit of log-normalized distribution of EW for SCUBA2 detections is $\mu(\log \text{EW}/\text{\AA}) = 1.337$ and $\sigma(\log \text{EW}/\text{\AA}) = 0.321$ (red dashed line); while that for non-detections is $\mu(\log \text{EW}/\text{\AA}) = 1.784$ and $\sigma(\log \text{EW}/\text{\AA}) = 0.334$ (blue dashed line). The average value of non-detections is 60.81 \AA , which is close to that of the normal SDSS quasars at lower redshifts (Diamond-Stanic et al. 2009).

We then performed a K-S test to check the probability that the detections and non-detections are drawn from the same distribution. The p -value is 0.017, i.e. a >98% probability that the detections has a different distribution from the non-detections.

The L_{FIR} was integrated using SCUBA2 850 μm detections with the modeled FIR SEDs described in Section 4.1. The SCUBA2 850 μm band has a rest wavelength coverage of $\sim 113\text{--}130\mu\text{m}$ corresponding to redshift of 5.5–6.5, which represents the dust emission from the host galaxy. On the other hand, the bolometric luminosity derived from optical band is dominated by the central quasar. Here we plot the relation between luminosity and equivalent width in Figure 9. The gray points are BAL quasars excluded in the following analysis. We fit detections using linear regression with the expectation maximization algorithm in the IRAF STS-

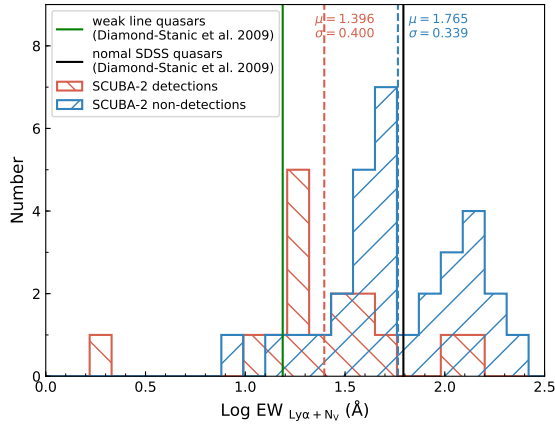


Figure 8. Distribution of rest-frame EW ($\text{Ly}\alpha + \text{N v}$) for SCUBA2-detected (red) and non-detected (blue) quasars at $z \sim 6$. The red and blue dashed lines represent the mean values for SCUBA2 detections and non-detections, respectively. The green solid line shows the definition of WLQ with $\text{Ly}\alpha + \text{N v}$ rest-frame equivalent widths of $\text{EW} < 15.4 \text{ \AA}$, while the black solid line is the mean value of 62 \AA for the normal SDSS quasars. (e.g., Fan et al. 2006; Anderson et al. 2001; Collinge et al. 2005; Diamond-Stanic et al. 2009)

DAS package⁷ (Isobe et al. 1986). The best-fit for detections (filled symbols) are represented as black solid lines. We adopt the EM linear regression method taking into account the censored data (Isobe et al. 1986). The black dashed line is the best fitting. The results are as follows:

All sample :

$$\log \frac{L_{\text{FIR}}}{L_{\text{bol}}} = (-1.63 \pm 0.68) \log \text{EW} + (-0.32 \pm 0.76)$$

$$\log L_{\text{FIR}} = (-3.64 \pm 1.03) \log \text{EW} + (47.16 \pm 13.72)$$

Detection :

$$\log \frac{L_{\text{FIR}}}{L_{\text{bol}}} = (-26.81 \pm 165.46) \log \text{EW} + (-17.20 \pm 115.13)$$

$$\log L_{\text{FIR}} = (-9.17 \pm 5.96) \log \text{EW} + (117.66 \pm 77.52)$$
(4)

The EW is decreasing as the increasing of the ratio of L_{FIR} to L_{bol} (see Figure 9 middle & right). The Pearson correlation coefficient r -value of EW and L_{FIR} for the detected sample is -0.385 with p -value = 0.141; while the coefficient value of EW and L_{bol} r -value = -0.063 with p -value = 0.817. The correlation test does not suggest a strong correlation between EW and quasar bolometric

luminosity (also seen in Figure 9 left). The correlations here may suggest some intrinsic connection between UV emission line properties and FIR luminosity. This should be checked with larger sample in a wider range of FIR luminosities.

5.1.2. WLQs are not redder than normal ones

As we see in Section 5.1.1, the millimetre detections tend to have weaker $\text{Ly}\alpha$ emission. If this is due to the dust extinction, we may expect that the continuum will be also obscured and the average slope of optical continuum β for FIR detected quasars should be larger than that for FIR non-detections. Unfortunately, most of the spectra in this paper do not result in a good fit of the continuum due to low S/N and short wavelength coverage. Alternatively, we compare the colors of the FIR detected quasars and non-detections in different redshift bins, to test whether the dust in the host galaxies, traced by the FIR luminosity, obscure the central AGN and result in the weak broad line feature.

Figure 10 shows the y_{P1} band – $W1$ band color vs. redshift. If the dust can obscure the broad emission line region, the continuum will be also obscured and redder. The lines show the average $y_{\text{P1}} - W1$ color (rest-frame 1400–5000 \AA) with redshift bins of $\Delta z = 0.3$. There is no strong indication that the FIR detected sources have a redder color with respect to the non-detected ones. The K-S test for these subsample shows p -value of 0.959, i.e., a 95.9% probability that the colors of the detections and non-detections have the same distribution statistically. Moreover, if the dust of the host galaxy obscures not only the broad emission lines but also the continuum, the ratios of obscuration should be same at a given wavelength, so that the rest-frame equivalent width will not change.

5.2. Possible explanations for weak line feature

One possible explanation for weak line quasars is the ‘shielding-gas scenario’, firstly proposed by Wu et al. (2011). The shielding gas, located between the accretion disk and broad line region, blocks the nuclear ionizing continuum reaching the broad-line region (BLR), resulting in the observed weak line emission (Luo et al. 2015). They proposed that this shielding gas is the puffed accretion disk when the accretion rate is very high. A high Eddington ratio is a common property in most high redshift quasars, for instance, J0100+2802 in our SCUBA2 survey, which is a WLQ with $\text{EW}(\text{Ly}\alpha) < 10 \text{ \AA}$ and has a high Eddington ratio of ($L_{\text{bol}}/L_{\text{edd}} \sim 1$). It satisfies the scenario of ‘shielding-gas’, in the case ($L_{\text{bol}}/L_{\text{edd}} > 0.3$) the slim disk may have a geometrically thick inner region (Luo et al. 2015; Wang et al. 2014). Another possible explanation is the ‘evolution

⁷ <http://stdas.stsci.edu/cgi-bin/gethelp.cgi?emmethod.hlp>

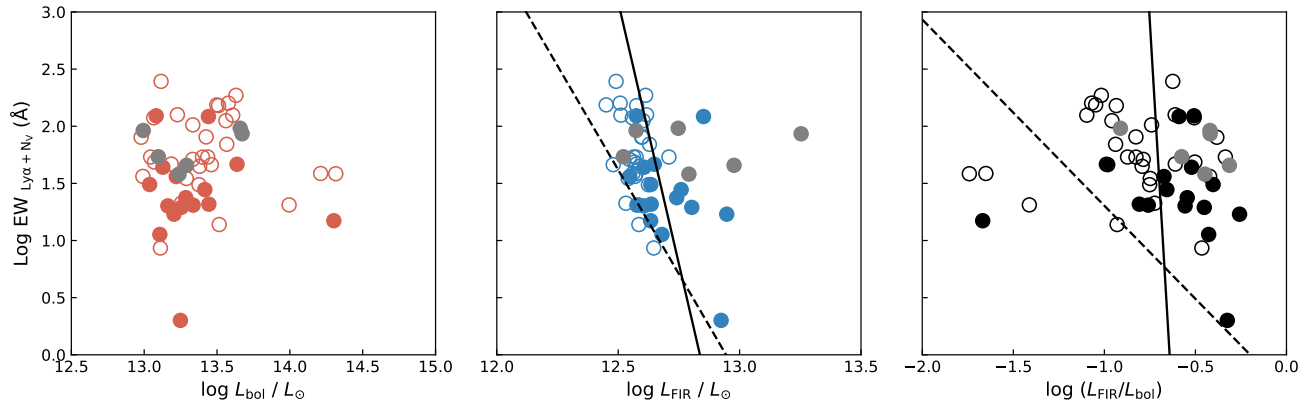


Figure 9. The relation between the equivalent width and (a) AGN bolometric luminosity, (b) FIR luminosity and (c) FIR-to-AGN luminosity ratio. The gray symbols are broad absorption line quasars, which have been excluded from the analysis. The filled and open symbols represent the FIR detections and 3σ upper-limit. The solid lines represent the linear fitting only for detections; and dashed lines are considered censored data using the survival analysis.

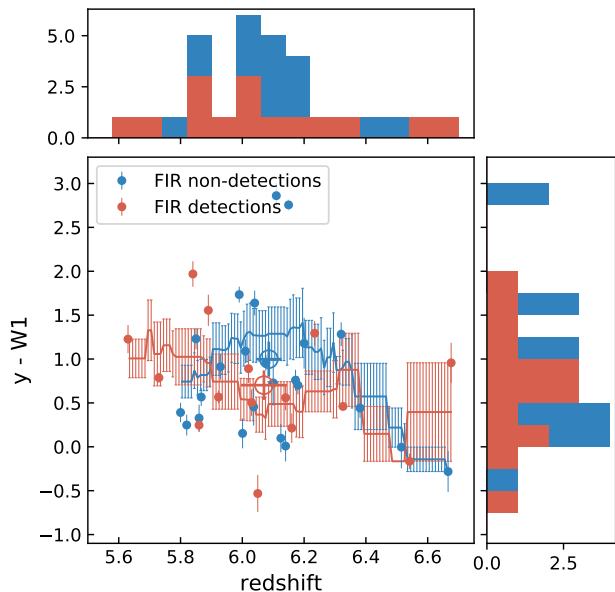


Figure 10. $y_{P1} - W1$ color in sub-samples of FIR detections (red) and non-detections (blue) at $z \sim 6$. The red (FIR detections) and blue (FIR non-detections) lines imply the average color in different redshift bins with width of $\Delta z = 0.3$. The large open symbols represent the average values of the sub-sample; the error bar is its standard error.

scenario’, at which BLR properties are always unusual, such as a low covering factor, an anisotropic ionizing source and so on (e.g., Plotkin et al. 2010; Hryniewicz et al. 2010; Laor & Davis 2011). In this case, the WLQ class represents an evolutionary stage, with a slow development of the BLR to manifest weak line phenomenon (Hryniewicz et al. 2010).

If these $z \sim 6$ WLQs are young AGNs evolve from galaxy mergers, it is natural that their host galaxies are actively forming stars with bright FIR luminosities. In the young AGN, the BLR is starting to develop slowly and/or the central quasar has some unusual accretion, SED and geometry. Direct evidences from observations to test these scenarios are still required. The physical mechanism could be different for individual WLQ at $z \sim 6$. More sub-mm/mm observations with NOEMA or ALMA for these $z \sim 6$ quasars, and further higher resolution and multi-wavelength observations (e.g., X-ray, radio) for these WLQs would benefit the study of co-evolution between the central AGN and its host galaxy.

6. CONCLUSION

In this paper we present JCMT SCUBA2 850 and $450 \mu\text{m}$ observations of 54 optical bright quasars with a wide range of quasar luminosities at $z \sim 6$. We construct a statistical sample to probe the far-infrared properties from the quasar host galaxies at the earliest epoch and study the evolution of quasars with redshift.

We concluded the following:

- We observed 54 quasars with an average $850 \mu\text{m}$ rms of $1.2 \text{ mJy beam}^{-1}$, and obtained detections for 20 sources ($> 4 \text{ mJy}$, at $> 3\sigma$). The new SCUBA2 detections have a wide flux range in $850 \mu\text{m}$ band of $3.34\text{--}16.85 \text{ mJy}$, and indicate FIR luminosities of 3.5×10^{12} to $3.0 \times 10^{13} L_{\odot}$, assumed a graybody SED. The stacked average flux density of $850 \mu\text{m}$ detections in our survey are $F_{s,850 \mu\text{m}} = 5.00 \pm 0.27 \text{ mJy}$, $\text{SNR} = 18.5$. For all $z \sim 6$ quasars, the value $F_{s,850 \mu\text{m}} = 2.30 \pm 0.16 \text{ mJy}$, $\text{SNR} = 14.4$. In our survey, P215–16 ($z = 5.73$) is the most luminous quasar at sub-mm band discovered at

$z \sim 6$ till now, with $F_{s,850\mu\text{m}} = 16.85 \pm 1.10 \text{ mJy}$ (SNR = 15.3) and $F_{s,450\mu\text{m}} = 26.03 \pm 7.78 \text{ mJy}$ (SNR = 3.3).

- In the individual SED fitting for SCUBA2 detections, the results imply extreme star formation rate in the range of 90 to $1060 M_{\odot} \text{ yr}^{-1}$ in the quasar host galaxies. The derived dust mass is in the range of $2.0\text{--}10.2 \times 10^8 M_{\odot}$. The AGN bolometric luminosities, estimated from M_{1450} , are in the range of 1.4×10^{13} to $4.7 \times 10^{14} L_{\odot}$; implying a black hole accretion rate in $9.6\text{--}318.5 M_{\odot} \text{ yr}^{-1}$ assuming the efficiency $\eta = 0.1$.
- The resulting median broad band SED for $z \sim 6$ quasars is similar to that at lower redshift, which indicates there is probably no evolution of quasar’s broad-band continuum emission properties with redshift.
- Luminous quasars are more rare at high redshift, e.g., $L_{\text{FIR}} > 10^{13}$; and FIR luminosity tend to be lower at $z \sim 6$ than lower redshift for a fixed bolometric luminosity, which may suggest a potential evolution of L_{FIR} to L_{bol} with redshift. However, this result is effected by the FIR detection limits and the selection effect of L_{bol} distribution.
- We measured the EW of the blended Ly α and N v emission lines for all samples, following the procedure of Diamond-Stanic et al. (2009). The WLQs’ distribution at high redshift has a lower peak and a larger dispersion, consistent with the results of Bañados et al. (2016).
- The EW (Ly α + N v) measurements show the high redshift sub-mm detected quasars tend to have the weaker emission line features. The p -value of K-S

test is 0.017, i.e., a $>98\%$ probability that the detections has a different distribution from the non-detections in statistic, which is also suggested in some previous work (e.g. Wang et al. 2008b).

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APPENDIX

A. MEASURE EW (Ly α + N v) FROM NIR SPECTRUM

The NIR spectra of all 54 SHERRY sample are shown here. These spectra include 52 published spectra and 2 unpublished spectra. For each individual spectrum, we fit a power law of the form $f_{\lambda} = C \times \lambda^{\beta}$ to continuum regions uncontaminated by emission lines following the procedure in Diamond-Stanic et al. (2009), shown in blue region of Figure 11. The fitted continuum is shown as blue dashed line. The fitted slope β and the derived EW (Ly α + N v) shown in the right top corner. We note that these spectra often have low S/N or absorption line features, which introduce many uncertainties to measure their EW (Ly α + N v).

A.1. Estimation of EW (Ly α + N v)

In order to estimate EW (Ly α + N v) of our quasar sample as precisely as possible and compare with previous work, we followed the procedure of Diamond-Stanic et al. (2009). For each individual spectrum, we fit a power law of the form $f_{\lambda} = C \times \lambda^{\beta}$ to continuum regions uncontaminated by emission lines (1285–1295, 1315–1325, 1340–1375, 1425–1470, 1680–1710, 1975–2050 and 2150–2250 Å) (Diamond-Stanic et al. 2009). To obtain a good fit for continuum, we first fitted the power law index β , where β in $[-2, -1]$. In many cases, a good fit is not possible due to the low S/N

spectrum and broad absorption feature so we needed to fix β . Here we assumed that the slope is the average quasar UV continuum slope $\beta = -1.5$, which is the most common value (e.g., Vanden Berk et al. 2001). After subtracting the power-law continuum from the spectrum, we then determined EW by simply integrating the flux above the continuum between $\lambda_{\text{rest}} = 1160 \text{ \AA}$ and $\lambda_{\text{rest}} = 1290 \text{ \AA}$. It is a blended emission feature dominated by the Ly α and N v, includes Ly α $\lambda 1216$, N v $\lambda 1240$, and Si II $\lambda 1263$ emission lines (Diamond-Stanic et al. 2009). Finally, we obtained EW of Ly α + N v shown in Table 5. Note these may be the lower limits because at these high redshifts some fraction of the Ly α line is being absorbed by the neutral IGM.

A.2. Uncertainties of measurement

These spectra are all discovery spectra, and some of them have low S/N due to short exposure time. Except J0305–3150, J2348–3054, J10480109, P167–13 and P338+29, the spectra have a narrow wavelength coverage with $\lambda_{\text{rest}} < 800 \text{ \AA}$, i.e., they just cover the nearby regions of Ly α line. Furthermore, as can be seen from Figure 11, absorption line features are universal phenomena in high redshift quasars. These are all difficulties to determine the continuum for most of spectra, and cause many uncertainties to measure their EW (Ly α + N v). The table 5 also lists the EW of Ly α + N v and the quasar type from literatures. In this paper, the typical uncertainties only from the fitting is about $\sim 5 \text{ \AA}$. But there are some other uncertainties, such as low spectral qualities, continuum etc. These will affect the EW estimate; thus, more accurate measurement is necessary for the future high quality spectra.

A.3. Comparisons to previous literature measurements

The EW often more or less differs from the values in literatures due to the different measure methods. Here we discuss the differences between this work and the literature results for each targets.

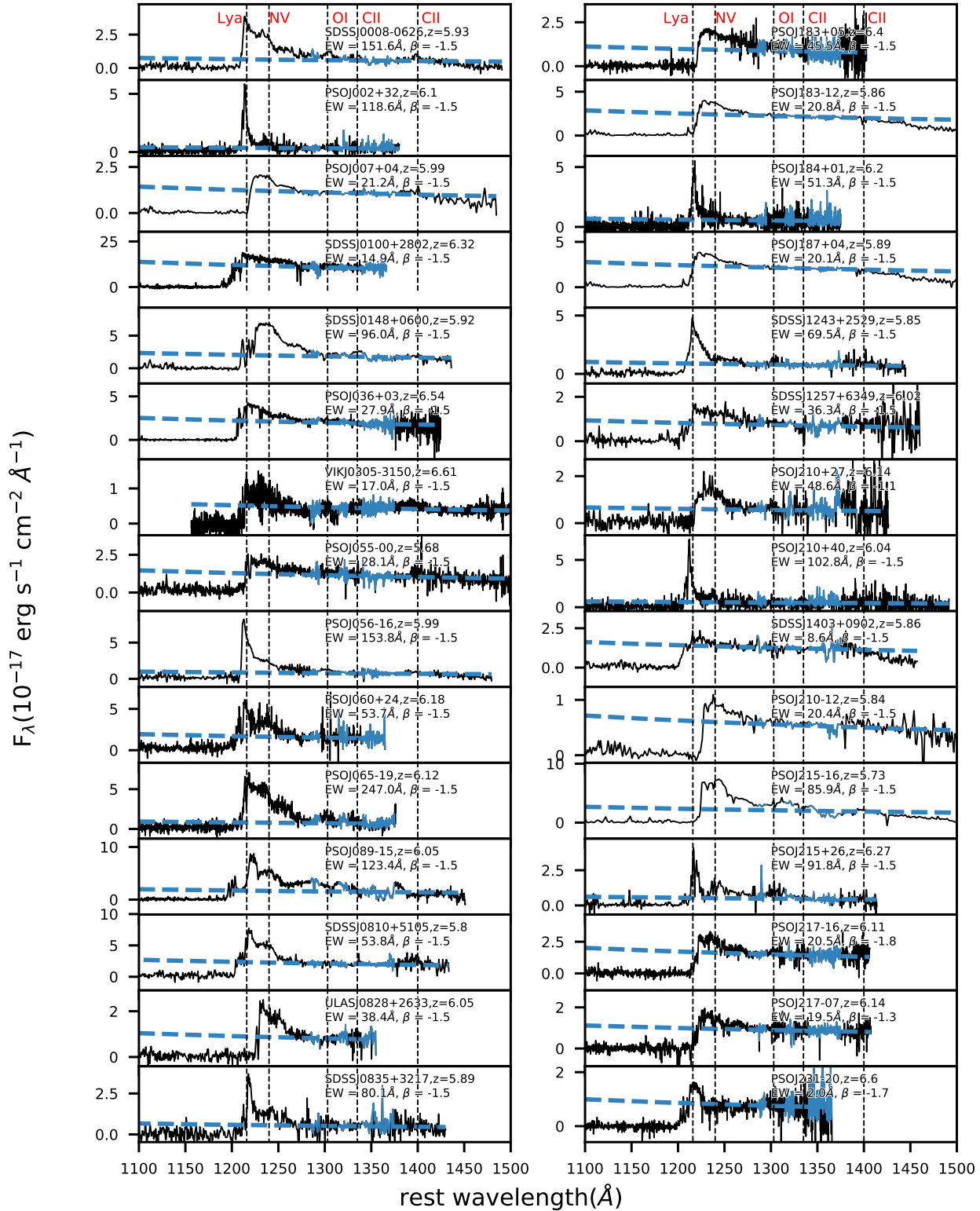
J0008-0626, J0148+0600, J0842+1218, J0850+3246, J1207+0630, J1257+6349, J1403+0902: Jiang et al. (2015) selected the optical spectra only cover a short wavelength range in the rest frame ($100 \sim 150 \text{ \AA}$), and this range contains several UV emission lines, so they were not able to reliably measure the slopes of the continua. And then they measured EW by fitting a half Gaussian profile (Ly α) and a full Gaussian profile (N v) simultaneously, for some quasars with prominent Ly α emission. For the quasars without prominent Ly α emission (i.e. J1403+0902, EW (Ly α + N v)=8 \AA), they simply integrated flux (instead of Gaussian fitting) above the continuum, which is the same method with this work and has the similar value of 8.6 \AA in this paper.

J0100+2802: Wu et al. (2015) reported the EW of the Ly α + N v emission lines roughly measured from the LBT spectrum is about 10 \AA . In this work, our value is about 14.9 \AA , which is within the uncertainty of measurement.

P183-12, P210-12: Bañados et al. (2014) suggested P183-12 and P210-12 belong to the weak-line quasar classification with EW (Ly α + N v) of 11.8 and 10.7 \AA respectively. They noted that the EW is very dependent on the continuum fit estimate. And since most of the spectra do not cover the region with $\lambda_{\text{rest}} > 1500 \text{ \AA}$, a good fit to the continuum is challenging. Thus, the uncertainties in their EW estimates are of the order of 25%. In this work, we measured EW (Ly α + N v) of 20.8 and 20.4 \AA respectively. The typical uncertainties only from the fitting is $\sim 5 \text{ \AA}$. Considering some other uncertainties, such as low spectral qualities, the uncertainties should be larger ($> 5 \text{ \AA}$) and our results have some differences with Bañados et al. (2014), but still within the uncertainties of measurement.

P215-16: Morganson et al. (2012) fitted a continuum with a $\lambda^{-3.05}$ power law and fitted the Ly α + N v using a single Gaussian, assuming it is a non-BAL. The derived EW is $109.5 \pm 83.1 \text{ \AA}$. In this paper, our result is about 85.9 \AA within the uncertainties of measurement.

In summary, our EW measurements in most cases are consistent with that in the literature, and the differences are mainly due to a different assumptions and/or low S/N of the data.

Figure 11. The NIR spectra of $z \sim 6$ quasars in SCUBA2 survey

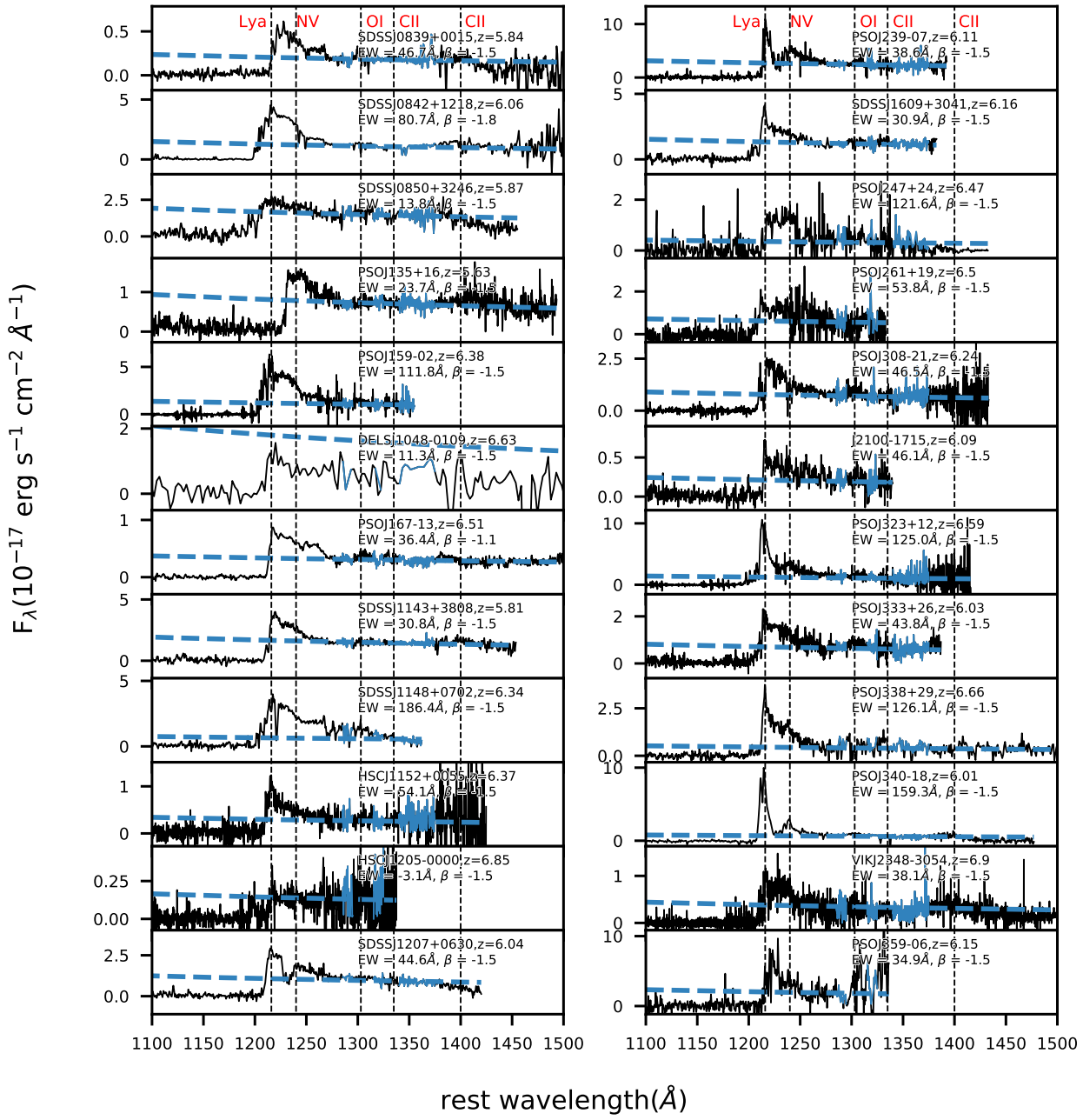


Figure 12. -Continued

Table 6. NIR Spectroscopic Observations

| source name | Telescope/Instrument | spectral coverage | spectral resolution R | Ref. |
|-------------|---|-------------------|-----------------------|--------|
| | | μm | | |
| (1) | (2) | (3) | (4) | (5) |
| J0008-0626 | MMT/Red Channel | 0.67 - 1.03 | 800 | (1) |
| P002+32 | P200/DBSP | 0.50 - 0.98 | 1500 | (2) |
| P007+04 | VLT/FORS2 | 0.71 - 1.04 | 1400 | (2) |
| J0100+2802 | Magellan/FIRE | 0.65 - 1.00 | 6000 | (3) |
| J0148+0600 | MMT/Red Channel | 0.67 - 0.99 | 800 | (1) |
| P036+03 | NTT/EFOSC2, Magellan/FIRE, Keck I/LRIS, VLT/FORS2 | 0.75 - 1.07 | 6000 | (2) |
| J0305-3150 | Magellan/FIRE | 0.88 - 2.27 | 6000 | (4) |
| P055-00 | VLT/FORS2 | 0.71 - 1.04 | 1400 | (2) |
| P056-16 | MMT/Red Channel | 0.67 - 1.03 | 800 | (2) |
| P060+24 | P200/DBSP | 0.50 - 0.98 | 1500 | (2) |
| P065-19 | P200/DBSP | 0.50 - 0.98 | 1500 | (2) |
| P089-15 | Keck I/LRIS | 0.54 - 1.02 | 1500 | (2) |
| J0810+5105 | MMT/Red Channel | 0.70 - 0.97 | 800 | (8) |
| J0828+2633 | Gemini/GMOS | 0.53 - 0.96 | 1300 | (5) |
| J0835+3217 | MMT/Red Channel | 0.72 - 0.99 | 800 | (8) |
| J0839+0015 | VLT/FORS2 | 0.71 - 1.04 | 1390 | (6) |
| J0842+1218 | Keck/ESI | 0.39 - 1.09 | 1500 | (1) |
| J0850+3246 | MMT/Red Channel | 0.68 - 1.00 | 800 | (1) |
| P135+16 | LBT/MODS, VLT/FORS2 | 0.71 - 0.99 | 1400 | (2) |
| P159-02 | VLT/FORS2 | 0.75 - 1.00 | 1400 | (2) |
| J1048-0109 | Magellan/FIRE | 0.76 - 2.36 | 300-500 | (7) |
| P167-13 | VLT/FORS2, Magellan/FIRE | 0.71 - 2.43 | 1400 | (2) |
| J1143+3808 | MMT/Red Channel | 0.70 - 0.99 | 800 | (2)(8) |
| J1148+0702 | Magellan/FIRE | 0.68 - 1.00 | 1500 | (8) |
| J1152+0055 | Subaru/FOCAS | 0.75 - 1.05 | 1200 | (9) |
| J1205-0000 | Subaru/FOCAS | 0.75 - 1.05 | 1200 | (10) |
| J1207+0630 | MMT/Red Channel | 0.68 - 1.00 | 800 | (1) |
| J183+05 | VLT/FORS2 | 0.71 - 1.04 | 1400 | (11) |
| P183-12 | VLT/FORS2, Magellan/FIRE | 0.71 - 1.04 | 1400 | (14) |
| P184+01 | LBT/MODS | 0.55 - 0.99 | 5000 | (2) |
| P187+04 | VLT/FORS2 | 0.71 - 1.04 | 1400 | (2) |
| J1243+2529 | MMT/Red Channel | 0.70 - 0.99 | 800 | (8) |
| J1257+6349 | MMT/Red Channel | 0.66 - 1.03 | 800 | (1) |
| P210+27 | MMT/Red Channel | 0.69 - 1.02 | 800 | (2) |
| P210+40 | P200/DBSP | 0.48 - 1.05 | 1500 | (2) |
| J1403+0902 | MMT/Red Channel | 0.68 - 1.00 | 800 | (1) |
| P210-12 | VLT/FORS2 | 0.71 - 1.04 | 1400 | (14) |
| P215-16 | MMT/Red Channel | 0.50 - 1.13 | 800 | (15) |
| P215+26 | VLT/FORS2 | 0.68 - 1.03 | 1400 | (13) |
| P217-16 | VLT/FORS2 | 0.71 - 1.00 | 1400 | (2) |
| P217-07 | Magellan/LDSS3 | 0.65 - 1.01 | 4000 | (2) |
| P231-20 | Magellan/FIRE, VLT/FORS2 | 0.71 - 1.04 | 1400 | (11) |
| P239-07 | MMT/Red Channel, VLT/FORS2 | 0.71 - 0.99 | 1400 | (2) |
| J1609+3041 | MMT/Red Channel | 0.70 - 0.99 | 800 | (8) |
| P247+24 | Magellan/FIRE, VLT/FORS2 | 0.72 - 1.07 | 1400 | (11) |
| P261+19 | P200/DBSP | 0.55 - 1.00 | 1500 | (11) |
| P308-21 | VLT/FORS2 | 0.71 - 1.04 | 1400 | (2) |
| J2100-1715 | Gemini/GMOS | 0.65 - 0.95 | 1300 | (12) |
| P323+12 | VLT/FORS2, Magellan/FIRE | 0.75 - 1.07 | 1400 | (11) |
| P333+26 | Keck/LRIS | 0.60 - 0.97 | 5000 | (2) |
| P338+29 | MMT/Red Channel, Magellan/FIRE, LBT/MODS | 0.51 - 2.49 | 5000 | (2) |
| P340-18 | NTT/EFOSC2, LBT/MODS | 0.50 - 1.04 | 1400 | (2) |
| J2348-3054 | VLT/X-Shooter | 0.56 - 2.48 | 6000 | (6) |
| P359-06 | MMT/Red Channel | 0.47 - 0.96 | 800 | (2) |

NOTE— **References.** (1) Jiang et al. (2015); (2) Bañados et al. (2016); (3) Wu et al. (2015); (4) Venemans et al. (2013); (5) S. J. Warren et al. (in prep.); (6) Venemans et al. (2015a); (7) Wang et al. (2017); (8) Jiang et al. (2016); (9) Izumi et al. (2018); (10) Matsuoka et al. (2016); (11) Mazzucchelli et al. (2017); (12) Willott et al. (2010); (13) E. Bañados et al. (in prep.); (14) Bañados et al. (2014); (15) Morganson et al. (2012).

B. NIR SPECTROSCOPIC OBSERVATIONS FROM THE LITERATURE

Here we list the information of the near-infrared (NIR) spectra of $z \sim 6$ quasars in our survey from their discovery papers (e.g. Morganson et al. 2012; Venemans et al. 2013, 2015b,a; Bañados et al. 2014; Jiang et al. 2015; Bañados et al. 2016; Jiang et al. 2016; Matsuoka et al. 2016; Wang et al. 2016, 2017), shown in Table 6. And we also include some unpublished sources (S. J. Warren et al. in prep.; Bañados et al. in prep.).

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