

Detection of a coherent magnetic field in the Magellanic Bridge through Faraday rotation

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ABSTRACT

We present an investigation into the magnetism of the Magellanic Bridge, carried out through the observation of Faraday rotation towards 167 polarized extragalactic radio sources spanning the continuous frequency range of 1.3–3.1 GHz with the Australia Telescope Compact Array. Comparing measured Faraday depth values of sources ‘on’ and ‘off’ the Bridge, we find that the two populations are implicitly different. Assuming that this difference in populations is due to a coherent field in the Magellanic Bridge, the observed Faraday depths indicate a median line-of-sight coherent magnetic-field strength of $B_{\parallel} \simeq 0.3 \mu\text{G}$ directed uniformly away from us. Motivated by the varying magnitude of Faraday depths of sources on the Bridge, we speculate that the coherent field observed in the Bridge is a consequence of the coherent magnetic fields from the Large and Small Magellanic Clouds being pulled into the tidal feature. This is the first observation of a coherent magnetic field spanning the entirety of the Magellanic Bridge and we argue that this is a direct probe of a ‘pan-Magellanic’ field.

Key words: polarization – Magellanic Clouds – galaxies: magnetic fields.

1 INTRODUCTION

The Large Magellanic Cloud (LMC) and Small Magellanic Cloud (SMC) is a highly studied galaxy pair. Due to their close proximity to the Milky Way (MW), the Magellanic Clouds allow astronomers to study galaxy interactions and evolution in unprecedented detail. The ongoing interaction between the galaxy pair, and possibly the MW, has led to the creation of the Magellanic Bridge (MB), the Magellanic Stream and the Leading Arm (see Besla et al. 2010 and D’Onghia & Fox 2016 for a complete review). Each of these tidal features can be identified through the presence of large amounts of H I gas. Most prominent of these features is perhaps the MB (Hindman, Kerr & McGee 1963) – a contiguous, gaseous tidal feature that spans the region between the LMC and SMC. We assume that the MB is located at a distance of 55 kpc, the mean distance to the LMC and SMC (Walker 1999). We also assume that the bulk of the H I emission in the MB has a radial velocity in the range $+100 \text{ km s}^{-1} \leq v_{\text{LSR}} \leq +300 \text{ km s}^{-1}$ (Muller et al. 2003; Putman et al. 2003). The tidal remnant is thought to have formed ~ 200 Myr

ago when the LMC and SMC were at their closest approach to one another (Gardiner & Noguchi 1996; Besla et al. 2012).

Tidal tails, streams and bridges play an important role in the evolution of the parent galaxies as well as the host environment, as they serve as a siphon for galactic material to be dispensed into the diffuse intergalactic medium. It can be posited that a pre-existing magnetic field could follow the movement of neutral gas into the intergalactic medium. The stretching and compressing of tidally stripped gas may then serve as a mechanism for the amplification of any existing magnetic fields (Kotarba et al. 2010). Thus, the stripping of tidal debris may be partially responsible for the distribution of magnetic fields over large volumes. What remains unclear is the importance and role of magnetic fields within tidal features.

The association between tidal remnants and magnetic fields has been studied for nearly two decades. Classically, the radio continuum tidal bridge connecting the ‘Taffy’ galaxies (Condon et al. 1993) was estimated as having a similar magnetic-field strength to the pre-collision galaxies and the field lines appeared to be stretching across the space between the galaxy pair. More recently, tidal dwarfs within the Leo Triplet and Stephan’s Quintet have been shown to possess coherent magnetic fields and have total magnetic-field strengths of $B_T = 3.3 \pm 0.5 \mu\text{G}$ and $B_T = 6.5 \pm 1.9 \mu\text{G}$, respectively (Nikiel-Wroczyński et al. 2013a,b).

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Decades of research using optical polarized starlight has shown that polarization vectors in the plane of the sky trace out a path from the SMC along the western Bridge oriented in the direction of the LMC (Mathewson & Ford 1970a,b; Schmidt 1970, 1976; Magalhaes et al. 1990; Wayte 1990; Lobo Gomes et al. 2015). Due to the limited number of stars with which one can carry out optical polarimetry studies, all previous claims of the existence of a coherent magnetic field spanning the entire Magellanic System have had to be speculative due to the lack of information stemming from the diffuse MB.

Studies of Faraday rotation of background polarized radio sources towards the LMC have determined that the galaxy has a coherent magnetic field of strength $\sim 1 \mu\text{G}$ (Gaensler et al. 2005). Mao et al. (2008) observed the SMC using both Faraday rotation measures (RMs) and polarized starlight. Through careful consideration of the Galactic foreground, they constructed 3D models for the magnetic field and showed that the orientation of the field has a possible alignment with the MB.

A similar investigation into Faraday rotation towards extragalactic polarized sightlines has shown that a high-velocity cloud (HVC) in the Leading Arm hosts a coherent magnetic field (McClure-Griffiths et al. 2010). In such an instance, a magnetic field could work to prolong the structural lifetime of the HVC as it is accreted on to the MW disc. While the exact origin of the magnetic field in this HVC remains unclear, it is plausible that the HVC fragmented from a magnetized Leading Arm. Therefore, the observed magnetic field in the HVC would be a consequence of the initial seed field followed by compression and amplification due to the MW halo.

Although magnetic fields have been found in the SMC, LMC and some HVCs, none of the previous investigations of magnetism in the Magellanic System have directly confirmed the existence of the *Pan-Magellanic Field* – a coherent magnetic field connecting the two Magellanic Clouds.

1.1 Faraday rotation

Complex linear polarization is an observable quantity and can be defined as

$$\mathcal{P} = Q + iU = p_0 e^{2i\Psi}, \quad (1)$$

where Q and U are the observed linearly polarized Stokes parameters, p_0 is the polarization fraction intrinsic to the source and Ψ is the observed polarization angle, also defined as

$$\Psi = \frac{1}{2} \arctan \frac{U}{Q}. \quad (2)$$

The polarization angle is rotated from its intrinsic value (Ψ_0) any time the emission passes through a magnetoionic material. This effect is known as Faraday rotation. The total observed Faraday rotation, defined $\Delta\Psi/\Delta\lambda^2$, is known as the RM.

When the rotating material is located along the line of sight, Faraday rotation can serve as a powerful tool to analyse magnetism. In the simple case of a thermal plasma threaded by a single magnetic field, the intrinsic polarization angle is rotated by $\Delta\Psi = \text{RM}\lambda^2$ rad. However, recent studies have shown that the RM may offer an incomplete, or misleading diagnostic of the actual polarization properties along the line of sight (O’Sullivan et al. 2012; Anderson, Gaensler & Feain 2016) and that many sources cannot be described by a single RM. It is therefore more robust to discuss the polarized signal in terms of its Faraday depth (ϕ), as first derived by Burn (1966). The Faraday depth encodes the electron density (n_e , in

cm^{-3}) and magnetic-field strength along the line of sight (B_{\parallel} , in μG) according to

$$\phi(L) = 0.812 \int_L^0 n_e B_{\parallel} dl, \quad (3)$$

where L is the distance through the magnetoionic material in parsecs. The sign of the Faraday depth is indicative of the orientation of the magnetic field with a positive ϕ signifying the field to be oriented towards the observer and a negative ϕ implying a field that is pointing away.

The measured ϕ_{obs} for an extragalactic source behind the MB is a summation of the various Faraday depth components along the line of sight and can be broken down into its constituent parts as follows:

$$\phi_{\text{obs}} = \phi_{\text{intrinsic}} + \phi_{\text{IGM}} + \phi_{\text{MB}} + \phi_{\text{MW}}, \quad (4)$$

where $\phi_{\text{intrinsic}}$ is the Faraday depth that is associated with the polarized emitting source, ϕ_{IGM} is any rotation due to the intergalactic medium, ϕ_{MB} is our targeted Faraday depth due to the posited MB magnetic field and ϕ_{MW} is the Faraday rotation due to the foreground MW. Although $\phi_{\text{intrinsic}}$, ϕ_{IGM} and ϕ_{MW} are present along all sightlines, ϕ_{MW} is likely to dominate the observed signal. This assumption appears to have been well justified in Taylor, Stil & Sunstrum (2009), whereby mapping the RMs of extragalactic polarized sources from the NRAO VLA Sky Survey (NVSS) revealed local structures in the Galaxy. Therefore, by observing polarized sources with sightlines that do not intersect the MB, we will be able to correct for the Galactic foreground, leaving the residual ϕ to represent the intrinsic properties of the background source and the MB contribution. The intrinsic polarized properties of each polarized source are random and considered to have a negligible effect on the overall statistics for a large sample.

If there exists a coherent magnetic field threading the MB, observations of linearly polarized background radio sources may hold the key to its discovery. In this work, we use detailed measurements of the Faraday depth of background, extragalactic polarized sources to investigate the existence of a coherent magnetic field spanning the MB. We describe our source selection process and observations in Section 2, followed by data reduction and processing in Section 3. We present our results in Section 4, which include the fitting and subtraction of the MW foreground. Section 5 motivates different distributions of ionized gas and the subsequently derived magnetic-field strengths. In Section 6, we discuss the possible origins and implications of the pan-Magellanic Field. A summary is presented in Section 7. Table A1 contains a glossary of the variables used in this paper.

2 OBSERVATIONS AND DATA

2.1 Source selection

For this investigation, we observed a subset of polarized sources that were originally identified through the reduction and re-processing of archival continuum data of the western MB (see Muller et al. 2003 for a summary of observations). In the literature, this region has been referred to usually as either the ‘Wing’ or ‘Tail’ (Brüns et al. 2005; Lehner et al. 2008), and we make reference to this region as the ‘Wing’, exclusively (See Fig. 1 for location). The H I observations of the ‘Wing’ had simultaneously observed the continuum emission associated with this region. The source-finding algorithm *Aegean* (Hancock et al. 2012) was used to identify polarized sources in the

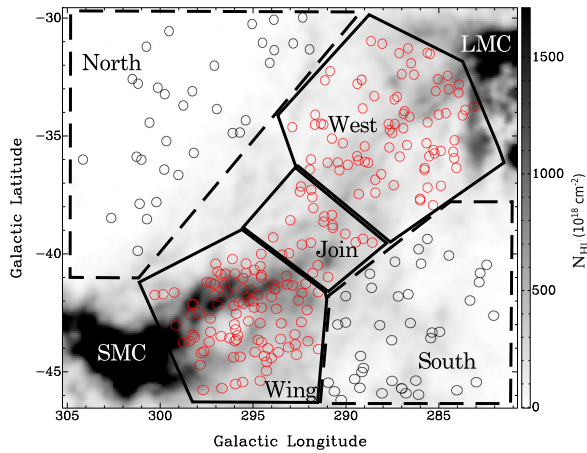


Figure 1. Neutral hydrogen column density for the velocity range of $+100 \leq v_{\text{LSR}} \leq +300 \text{ km s}^{-1}$ of the MB region from the GASS survey (McClure-Griffiths et al. 2009, Kalberla et al. 2010), overplotted with the positions of observed radio sources. Each pointing is associated with a region name denoted by the text in the enclosed areas. Red circles (pointings enclosed by a solid line) are sources where the MB is considered to intersect the background source’s line of sight, whereas sources marked by black circles (pointings enclosed by a dashed line) are considered as having lines of sight that are not contaminated by the MB. These latter sources were observed in order to subtract the Faraday rotation contribution from foregrounds and backgrounds.

final, deconvolved continuum images. From this original sample, we targeted 101 polarized sources for follow-up observations.

An additional 180 radio sources were targeted in order to extend the investigation across the entirety of the MB and surrounding area. Motivated by the changing morphology and kinematics of the Bridge, we separate these additional sources into regions ‘West’, ‘Join’, ‘North’ and ‘South’. These additional radio sources were selected from the Sydney University Molonglo Sky Survey (SUMSS; Mauch et al. 2003) as having a Stokes I flux $\geq 100 \text{ mJy}$ at 843 MHz for the region labelled ‘West’ and $\geq 150 \text{ mJy}$ for regions ‘Join’, ‘North’ and ‘South’. Fig. 1 gives a summary of the pointing regions observed overlaid on a map of neutral hydrogen (H I) of the region from the Galactic All Sky Survey (GASS; McClure-Griffiths et al. 2009; Kalberla et al. 2010).

2.2 Observations

Observations of the 281 radio sources were taken over 3 d with the Australia Telescope Compact Array under project C3043. Taking advantage of the instantaneous broad bandwidths of the Compact Array Broad-band Backend (CABB; Wilson et al. 2011), the observations spanned the continuous frequency range of 1100–

3100 MHz. Each pointing was observed as a series of snapshots in order to improve uv -coverage. Phase calibrators were observed at least every 40 min. The bandpass and flux calibrator PKS B1934-638 was observed on 2015 March 14 and 2015 April 30 and PKS B0823-500 was observed as the bandpass calibrator on 2016 June 11. Polarization leakage calibrations were carried out using the aforementioned primary calibrators. On average, each pointing was observed for a total of 3 min. Due to the nature of the source selection associated with the ‘Wing’ and the possibility that sources could be weak in total intensity, the initial 3 min of observation was sometimes not enough to reach a sufficient signal to noise. Additional observations were made as a single hour-angle uv -cut on 2016 June 11 in order to improve our sensitivity limits for points that were not bright enough in polarization nor total intensity to be confidently detected with our initial observations. A summary of the observations is listed in Table 1 and the representative uv -coverage for any source in each region is shown in Fig. 2.

3 DATA REDUCTION AND EXTRACTION

Observations were calibrated and imaged in the MIRIAD software package (Sault, Teuben & Wright 1995) using standard routines. Flagging of the data was done largely with the automated task PGFLAG, with minor manual flagging being carried out with tasks BLFLAG and UVFLAG. Naturally weighted Stokes I , Q , U and v maps were made using the entire 2 GHz bandwidth. Deconvolution of the multifrequency data set was performed on the dirty maps with the task MFCLEAN. Cleaning thresholds were set to be three times the rms Stokes V levels ($3\sigma_V$) for Stokes Q and U and $5\sigma_V$ for Stokes I . Images were convolved to a common resolution of 8 arcsec, which corresponds to a linear scale of 2 pc at the assumed distance to the MB of 55 kpc.

From the broad-band 2 GHz images, images of linearly polarized intensity (\mathcal{P}) were made with the task MATH. The total polarized flux of a target was extracted from an aperture 8 arcsec in diameter centred on the peak polarization pixel with noise estimates ($\sigma_{\mathcal{P}}$) measured as the rms residuals from a source-extracted image. A target was considered ‘polarized’ if the integrated polarized flux was greater than $8\sigma_{\mathcal{P}}$. This method of imaging will lead to bandwidth depolarization for sources with absolute Faraday depths greater than $\sim 90 \text{ rad m}^{-2}$; however, we consider the number of sources rejected due to high Faraday rotation to be negligible and has no impact on our final science goals.

Imaging with narrow bandwidths decreases the signal to noise in addition to reducing the resolution in Faraday depth space, while broad bandwidths decrease the maximum observable scale in Faraday space, as well as the maximum observable Faraday depth. In order to minimize the bandwidth depolarization and maintain a desirable signal-to-noise ratio, Stokes I , Q , U images were made

Table 1. Summary of the observations. Column 1 gives the array configuration; column 2 gives the regions targeted (as defined in Section 2); column 3 lists the length of the observing run and column 4 gives an approximation for the total integration time per source. Column 5 gives the UT date of the commencement of the observations.

Array config.	Obs. targets	Obs. length (h)	Time on-source (min)	Obs. date
6C	Wing, West	12	2.5	2015 Mar 14
6A	Wing, West	15	1.5	2015 Apr 30
6A	Join, North, South	15	3	2015 Apr 30
1.5B	Wing (subset)	3	5	2016 Jun 11

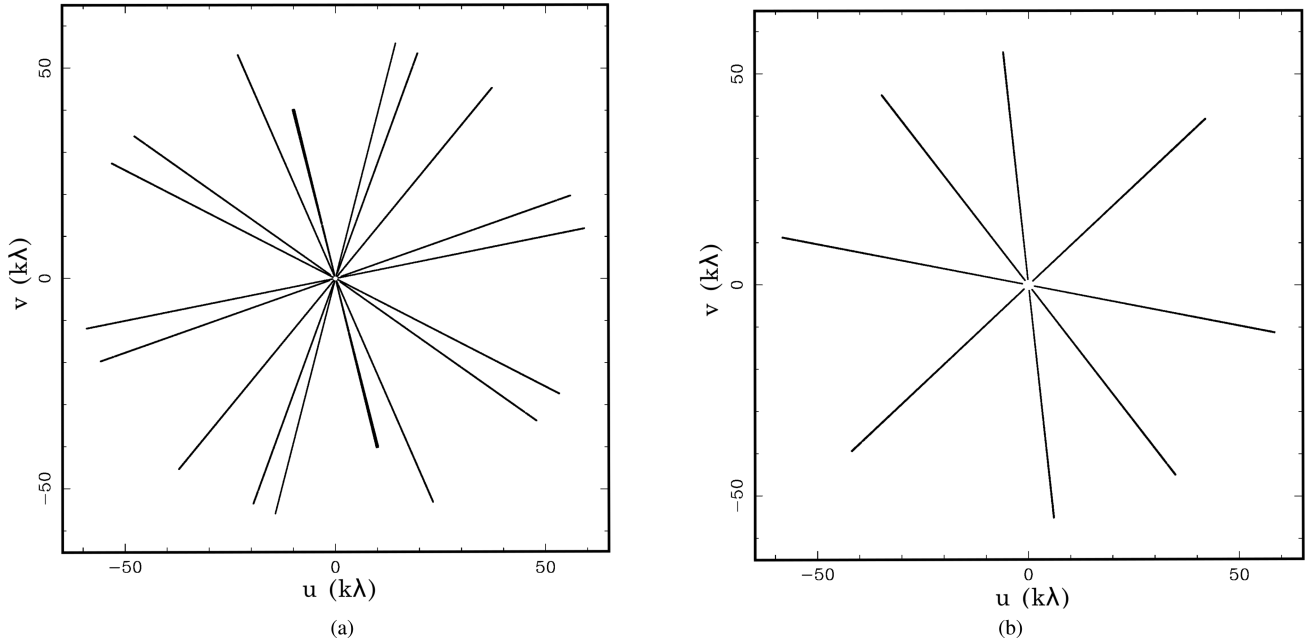


Figure 2. Typical uv -coverage of a single radio source associated with (a) the ‘Wing’ and ‘West’ and (b) ‘Join’, ‘North’, ‘South’.

Table 2. Summary of total number of points observed per region and total number of polarized sources accepted. In order to be accepted, a source must be detected to at least 8σ in the full bandwidth polarized intensity image. The ‘Wing’ region returns a higher fraction of polarized sources due to our previous knowledge of the polarization in this region.

Region	Observed	Polarized	Accepted fraction (per cent)
Wing	101	69	68
West	83	40	48
Join	23	15	65
North	34	22	65
South	40	21	53
	281	167	59

every 64 MHz – resulting in 27 channel maps spanning 1312–3060 MHz.

As with the broad-band \mathcal{P} images, integrated fluxes were extracted from each map from an equivalent beam area centred on the pixel corresponding to the peak in \mathcal{P} . Error measurements were estimated as the rms-noise level from images created from the residual of the Stokes maps after the source aperture was blanked. With the exception of sources associated with the ‘Wing’, all targets are expected to be bright in total intensity. A further 10σ cut-off was imposed, and extracted spectra with fewer than 10 channels were discarded.

The procedures described above result in 167 sources with spectra in I , Q and U . Table 2 has a summary of the fraction of sources accepted per region. The ‘Wing’ region has an advantage in returning a higher number of polarized sources due to our previous knowledge of the polarized detections in the region. However, our data extraction method rejected multiple targets in the ‘Wing’ region for falling below the sensitivity threshold. Fig. 3 gives two examples of total ((a) and (c)) and polarized intensity ((b) and (d)) detected from extragalactic radio sources.

3.1 qu -fitting and ϕ determination

We adopt the fractional notation such that $q = Q/I$ and $u = U/I$, where the observable polarized fraction can be expressed as

$$p = \sqrt{q^2 + u^2}. \quad (5)$$

In working with fractional Stokes parameters, the wavelength-dependent depolarization effects are decoupled from spectral index effects.

To create our fractional polarized spectra, the Q and U spectra are divided by a model fit to the Stokes I spectrum. This approach avoids creating non-Gaussian noise and the propagation of small-scale spectral errors that may be present in the Stokes I spectrum. Using a bootstrap approach with 10 000 iterations, we fit a second-order polynomial to the Stokes I spectrum of each polarized source and calculate the standard deviation of the resultant q and u values for each frequency channel. The total error is considered to be the standard deviation of the bootstrapped values of q and u added in quadrature to the measured noise from the cleaned Stokes Q and U maps. The bootstrap method is necessary to correctly propagate the uncertainty due to the fit and has the overall effect of increasing the magnitude of the errors from what can be measured from the Stokes maps.

In order to extract the observed Faraday depth from our polarized signal, we must motivate a polarization model for the MB environment. External Faraday dispersion (Burn 1966) can be used as a proxy to measure fluctuations in the free-electron density or magnetic-field strength. This model has been used in numerous past studies of the polarization of galaxies, galaxy groups and clusters (Gaensler et al. 2005; Laing et al. 2008). Without an observed continuum-emission component of the MB, a single-component external Faraday dispersion model serves as an appropriate approximation to the polarization signal associated with the MB.

Polarization of this form displays a decreasing polarization fraction as a function of λ^2 . This depolarization can be defined as p/p_0 , where p is the observed polarization. This effect is most evident towards long wavelengths. Due to the purely external dependence

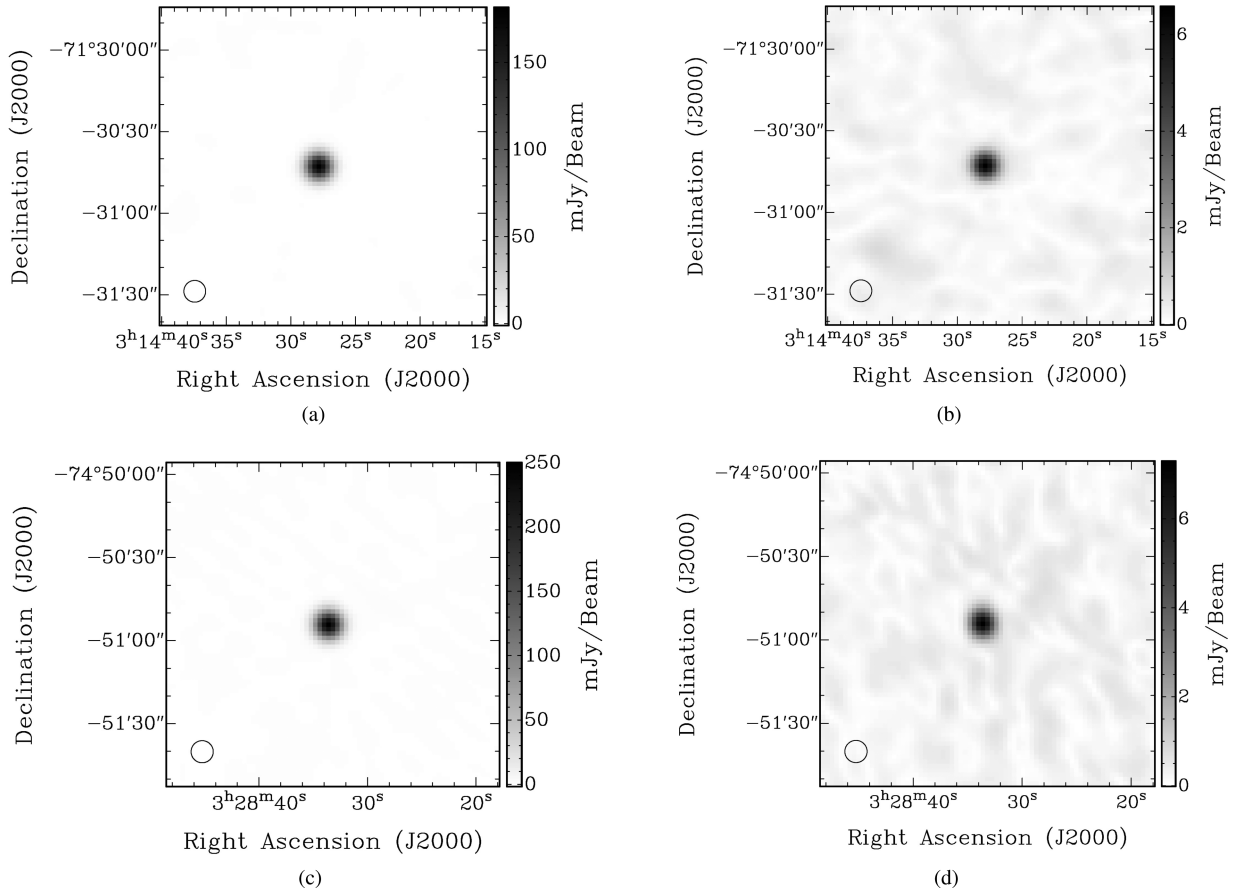


Figure 3. Example of two polarized sources detected in our survey: points ‘Join_08’ (a) and (b) and ‘West_02’ (c) and (d). Multifrequency images for total intensity (Stokes I) are shown in (a) and (c) and polarized intensity (\mathcal{P}) in (b) and (d). Both sources have been imaged using the full bandwidth available and the restoring beam is shown in the bottom left of each image.

of external Faraday dispersion, and its dependence on the size of the observing beam, this depolarization model is often referred to as ‘*beam depolarization*’. In this scenario, averaging the fluctuations across the entire beam area, the result is polarization of the form

$$\mathcal{P} = p_0 e^{2i(\Psi_0 + \phi_{\text{obs}}\lambda^2)} e^{-2\sigma_\phi^2\lambda^4}, \quad (6)$$

where ϕ_{obs} is the total observed Faraday depth value (equation 4) and σ_ϕ^2 characterizes the variance in Faraday depth on scales smaller than our beam.

We calculate the best-fitting ϕ_{obs} , σ_ϕ and Ψ_0 for each point source by fitting an external Faraday dispersion model (equation 6) simultaneously to the extracted $q(\lambda^2)$ and $u(\lambda^2)$ data. This technique is called *qu-fitting* and Sun et al. (2015) show it to be the best algorithm currently available for minimizing scatter in derived polarization parameters. We take a Monte Carlo Markov chain (MCMC) approach to fitting our complex polarization parameters by employing the ‘emcee’ PYTHON module (Foreman-Mackey et al. 2013). Compared to Levenburg–Marquardt fitting, MCMC better explores the parameter space, and returns numerically determined uncertainties for the model parameters. The log-likelihood of the complex polarization model of the joint $q - u$ chi-squared (χ^2) is minimized to find the best-fitting parameters. For each pointing, we initialize a set of 200 parallel samplers that individually and randomly explore the n -dimensional parameter space (where n is the degrees of freedom). Each of these samplers – called ‘walkers’ – iteratively cal-

culates the likelihood of a given location in parameter space and in doing so maps out a probability distribution for a set of parameters.

We initialize the walkers to random values of the free parameters and run three 300 iteration ‘burn-in’ phases where the samples settle on a parameter set of highest likelihood. The position history of the walkers is removed before initiating a 300-step exploration of the new parameter sub-space. The best-fitting model is calculated as the mean of the marginalized posterior distribution for each parameter. The parameter uncertainties are measured from the 1σ deviation of the walkers above and below the resultant best fit.

Fig. 4 gives an example solution from *qu-fitting*. The fractional Stokes spectra (p , q and u) versus λ^2 is shown in the top panel (a). Observed values are shown as black, blue and red points for p , q and u , respectively. The best-fitting solution is shown to trace the observed data. The best-fitting solution to Ψ versus λ^2 is given in the bottom panel (b). We attribute any deviation from the model to Faraday complexity of the source or a line-of-sight component that is not accounted for in the simple polarization model we assume (equation 6).

4 RESULTS

In addition to fitting the observed Faraday depth (ϕ_{obs}), our fitting routine also returns best-fitting values for all polarization parameters defined in equation (6), namely p_0 , Ψ_0 and σ_ϕ . A subsample of sources with the resultant best-fitting parameters is given in Table 3, with the full data set available in Appendix Table B1.

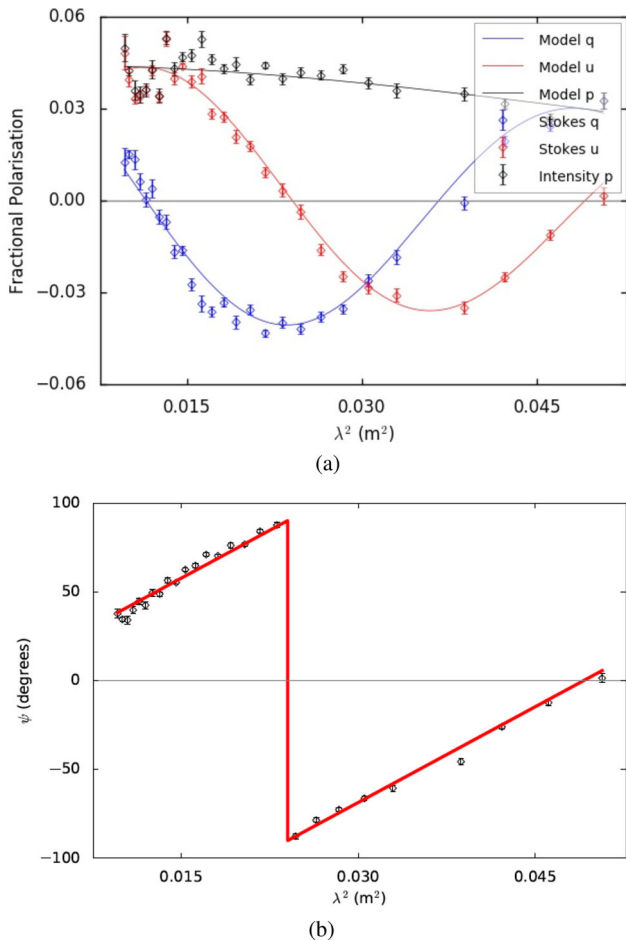


Figure 4. (a) Observed data and best-fitting solution for qu -fitting to a point in the ‘West’ region. Observed fractional Stokes q and u are shown as blue and red points, respectively, whereas the model solution is shown as blue and red lines. The observed and model polarized fraction is shown as black points and a black line for reference. (b) Corresponding fit to polarization angle (ψ) versus λ^2 for the aforementioned solution from qu -fitting.

Fig. 5 shows the best-fitting ϕ_{obs} of every polarized radio source plotted over the H I emission of the region from GASS (McClure-Griffiths et al. 2009; Kalberla et al. 2010). Red circles indicate a positive ϕ_{obs} and a field that is oriented towards the observer; blue circles, the opposite. Black crosses signify a ϕ_{obs} that is consistent with zero to $2 \times \Delta\phi$ where $\Delta\phi$ is the returned uncertainty in Faraday depth from qu -fitting.

We divide the observed polarized sources into two populations – those where the MB intersects the sightline to the polarized source and those with sightlines that are unaffected by the MB. We define an ‘on-Bridge’ region to be the area defined by a non-extinction corrected $\text{H}\alpha$ intensity of $I_{\text{H}\alpha} = 0.06 \text{ R}$, shown as the lowest contour in Fig. 8. The $\text{H}\alpha$ data set and subsequent analysis is discussed in more detail in Section 5.1. All sources associated with the ‘Wing,’ ‘West’ and ‘Join’ regions meet this criterion. The ‘North’ and ‘South’ regions are considered to be ‘off-Bridge’ and serve as a probe of the MW’s Faraday depth structure in the region.

Of all the ϕ_{obs} -values in the imaged region, 84 per cent are positive (red) and all of the negative (blue) and null (cross) Faraday depths are associated with the on-Bridge region (Fig. 5). Fig. 6 shows the ϕ_{obs} population of all on- and off-Bridge sources as a cumulative histogram and highlights the clear discrepancy in Faraday depths for each population. We test the statistical likelihood that the Faraday

depths associated with points on and off the Bridge come from a single population by performing a K-sample Anderson–Darling test on the best-fitting ϕ_{obs} -values for all sources that have been detected to $8\sigma_{\mathcal{P}}$ or higher in polarized intensity. The returned normalized test statistic allows us to reject the null hypothesis with a 99.992 per cent confidence level. The difference in Faraday depths between the populations of ϕ_{obs} -values indicates that the polarized radiation on and off the MB probes distinctly different magnetic environments.

4.1 Correcting for Faraday rotation due to the MW foreground

The amount of Faraday rotation observed towards an extragalactic point source (ϕ_{obs}) will always include some contribution from the MW. Therefore, before the line-of-sight magnetic-field strength can be estimated, the Galaxy’s contribution to the observed Faraday depth must be fit and corrected for. The 43 off-Bridge ϕ_{obs} can be described by a tilted-plane ϕ_{MW} -model, whose parameters are obtained using a non-linear least-squares fit to the data. The best-fitting solution was found to be of the form

$$\phi_{\text{MW}} = -0.511l + 1.28b + 225, \quad (7)$$

where l and b are the coordinates in Galactic longitude and latitude, respectively. The plane is shown in Fig. 7. By subtracting the resultant Faraday depth surface from all ϕ_{obs} , the residual Faraday depths (ϕ_{corr}) are considered to be foreground corrected.

We compare our MW Faraday depth model with similar models from Mao et al. (2008) and Oppermann et al. (2015). Testing a point in the centre of the ‘Join’ region ($l = 290^\circ$, $b = -38^\circ$), our fit returns a ϕ_{MB} -value of $+28 \text{ rad m}^{-2}$. At the same position, MW models from Mao et al. (2008) and Oppermann et al. (2015) return values of $+28 \text{ rad m}^{-2}$ and $+25 \text{ rad m}^{-2}$, respectively. The close agreement amongst all three MW models adds confidence to our MW correction.

We further test the validity of the foreground ϕ_{MW} -model by comparing the distributions of the uncorrected and corrected ϕ -values (ϕ_{obs} and ϕ_{MB} , respectively) for points in the ‘North’ and ‘South’ regions. If the assumptions made to create the foreground model were valid, the distribution of Faraday depths should become more similar after the foreground correction has been applied. We test this theory by conducting two separate Anderson–Darling tests on the ϕ_{obs} and ϕ_{corr} distributions for the two off-Bridge regions. We find that before the foreground correction is applied there is ~ 98 per cent confidence that the two background samples are drawn from different populations. Once our model is subtracted from the raw, observed Faraday depths, the likelihood that the two populations are unique drops to 67 per cent. At this level, there is no longer sufficient confidence to say they are not drawn from the same parent distribution. We therefore consider our simplified tilted-plane assumption of the Faraday depth distribution of the MW-foreground to be justifiable.

Fig. 8 shows the foreground-subtracted Faraday depths across the imaged region. We expect that after our foreground correction, the majority of off-Bridge sources would have values near zero, but this is not observed. We assume that the major cause for this discrepancy is that our foreground model is an oversimplification of the likely complex Faraday structure of the MW (Oppermann et al. 2015). We test the merit of a higher-order foreground Faraday depth model, but it produces minimal improvement while increasing the degrees of freedom. If our foreground fit was well founded, we would expect to have a mean ϕ_{corr} -value of off-Bridge points near zero: our sample returns $\bar{\phi}_{\text{off,corr}} = 0.3 \text{ rad m}^{-2}$ with a standard deviation of 12.0 rad m^{-2} , compared to $\bar{\phi}_{\text{off,obs}} = 25 \text{ rad m}^{-2}$ before

Table 3. A subset of measured and calculated source parameters. Columns (1) and (2) give the source location in Galactic longitude and latitude, respectively. Column (3) lists the integrated total intensity (I) over the full 2 GHz bandwidth with uncertainties. Integrated polarized flux (\mathcal{P}) with uncertainty is listed in column (4). Columns (5–8) give the best-fitting parameters returned from qu -fitting: namely, the intrinsic polarization fraction (column 5), the intrinsic polarization angle (column 6), the total Faraday depth along the line of sight (column 7) and the Faraday dispersion (column 8).

(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
l	b	I	\mathcal{P}	p_0	Ψ_0	ϕ_{obs}	σ_ϕ
($^\circ$)	($^\circ$)	(mJy)	(mJy)	(per cent)	($^\circ$)	rad m $^{-2}$	rad m $^{-2}$
291.778	−40.785	104.9 ± 0.2	3.2 ± 0.3	5.8 $^{+0.7}_{-0.6}$	37 $^{+4}_{-4}$	+6 $^{+3}_{-4}$	21 $^{+2}_{-2}$
288.589	−39.501	258.4 ± 0.07	3.9 ± 0.3	1.66 $^{+0.09}_{-0.08}$	82 $^{+3}_{-3}$	−0.2 $^{+2}_{-2}$	3 $^{+2}_{-2}$
290.958	−45.418	215 ± 3	7.1 ± 0.4	2.8 $^{+0.3}_{-0.3}$	49 $^{+5}_{-5}$	+13 $^{+4}_{-4}$	23 $^{+2}_{-2}$
285.625	−39.347	123.8 ± 0.2	1.6 ± 0.2	6.38 $^{+0.07}_{-0.07}$	143.9 $^{+0.5}_{-0.5}$	+26.8 $^{+0.3}_{-0.3}$	13.9 $^{+0.2}_{-0.2}$
296.659	−45.653	73.0 ± 0.2	6.8 ± 0.3	10.1 $^{+0.4}_{-0.4}$	71 $^{+2}_{-2}$	−13.1 $^{+0.9}_{-1}$	10.7 $^{+0.8}_{-0.8}$

subtracting the foreground. We note that the foreground ϕ_{MW} fit does not attempt to fit and subtract the Faraday rotation that is intrinsic to the background source. Schnitzeler (2010) estimates the spread in intrinsic Faraday depths of extragalactic sources to be $\simeq 6$ rad m $^{-2}$, which can account for much of the large standard deviation of the off-Bridge, foreground-corrected Faraday depths.

The uncertainty in the foreground Faraday depth subtraction must be included with the uncertainty in the Faraday depth of the on-Bridge sources. The magnitude of the increased error was determined through bootstrapping the foreground ϕ_{MW} surface 10 000 times with the standard deviation of the correction at each location ($\sigma_{\phi_{\text{MW}}}$). The mean uncertainty in Faraday depths through this method is $\bar{\sigma}_{\phi_{\text{MW}}} = 0.21$ rad m $^{-2}$. The expression for the total uncertainty in the Faraday depth of a background radio source therefore becomes

$$\Delta\phi(l, b)^2 = \Delta\phi_{\text{MCMC}}^2 + \sigma_{\phi_{\text{MW}}}^2(l, b), \quad (8)$$

where (l, b) are the coordinates of the point source.

We infer that the MB Faraday rotation, ϕ_{MB} , accounts for the majority of the residual rotation seen in points associated with the MB and assume for all further analysis that $(\phi_{\text{obs}} - \phi_{\text{MW}}) = \phi_{\text{corr}}$, where $\phi_{\text{corr}} \approx \phi_{\text{MB}}$. A map of foreground-corrected ϕ_{MB} is given in Fig. 8, which shows negative Faraday depths spanning the entirety of the MB. Analysis of this trend shows that 68 per cent of the polarized sources follow this trend to $2 \times \Delta\phi$, where $\Delta\phi$ is the calculated error in our Faraday depth measurement.

ϕ_{MB} may contain contributions from localized enhancements – such as H II and star formation regions – that may influence the observed magnetic field on scales to which we are sensitive (~ 2 pc). In order to identify any phenomenon that could influence the small-scale magnetic field fluctuations in the MB, we cross-reference our region of sky with Simbad (Wenger et al. 2000) and find seven molecular clouds (Chen et al. 2014) and four H II regions (Meaburn 1986; Bica et al. 2008) that are located in the ‘Wing’ region. Three of the molecular clouds and three H II regions are near the small patch of positive ϕ -values near $l = 295^\circ$, $b = -42^\circ$. These

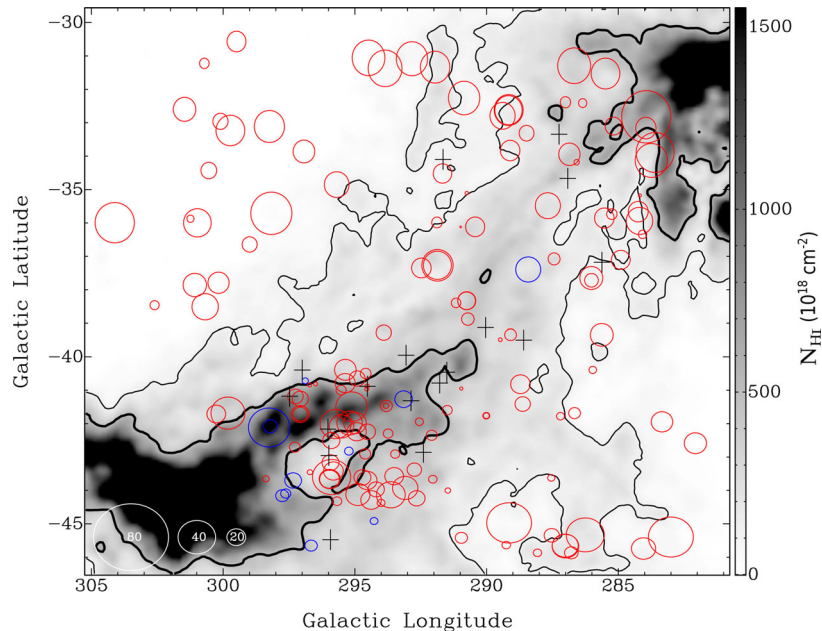


Figure 5. ϕ_{obs} values fit to an external Faraday dispersion model overlaid on a map of H I intensity from GASS (Kalberla et al. 2010) in the velocity range of $+100 \leq v_{\text{LSR}} \leq +300$ km s $^{-1}$. Black contours represent H I emissivity of 1.2 and 5.0×10^{20} cm $^{-2}$. The size of each circle is representative of the magnitude of ϕ , with scale-circles shown in the bottom-left corner. Red circles represent a line-of-sight magnetic field pointing towards the observer (positive ϕ), and blue circles show a field that is pointing away (negative ϕ). Black crosses show ϕ values consistent with zero to $2 \times \Delta\phi$.

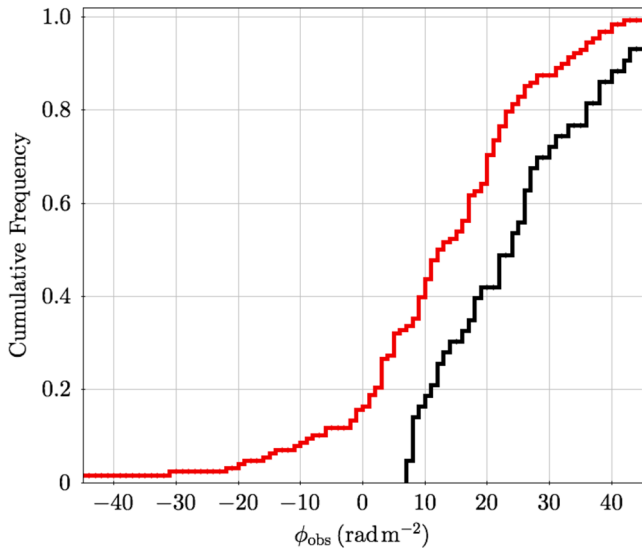


Figure 6. Cumulative histogram of ϕ_{obs} values for on-Bridge (red) and off-Bridge (black) sources. The figure is truncated at $\phi = \pm 45 \text{ rad m}^{-2}$ for clarity.

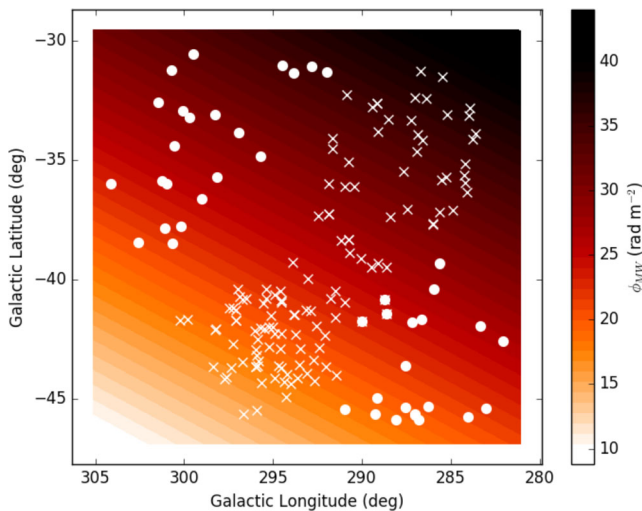


Figure 7. An estimation of the foreground- and background- ϕ covering our field of view, assuming the Faraday depth varies as a tilted plane across our imaged region. The fit used the 43 off-Bridge sources that are shown as white dots. The location of the on-Bridge sources are shown as white crosses.

individual molecular clouds do not directly align with any of the background sources at our physical-scale sensitivity of 8 arcsec.

5 THE LINE-OF-SIGHT MAGNETIC-FIELD STRENGTH

5.1 Emission measures

Our objective is to calculate the line-of-sight magnetic field (B_{\parallel}) associated with the MB; however, B_{\parallel} is degenerate with estimates of electron density (n_e). Therefore, an independent estimate of n_e is required. By making some assumptions about the line-of-sight depth of the ionized medium, it is possible to use observed $\text{H}\alpha$ intensities as a means to independently estimate n_e^2 by taking advantage of the implied emission measure (EM). The EM is defined as the integral

of the square of the electron density along the pathlength of ionized gas ($L_{\text{H II}}$) and can be derived from the measured $\text{H}\alpha$ intensity ($I_{\text{H}\alpha}$) in Rayleighs (R)¹

$$\text{EM} = \int_0^L n_e(l)^2 dl = 2.75 T_4^{0.92} I_{\text{H}\alpha} \text{ pc cm}^{-6}. \quad (9)$$

We utilize the work carried out by Barger, Haffner & Bland-Hawthorn (2013), which offers kinematically resolved intensities of the $\text{H}\alpha$ emission across our entire MB. The observations used in Barger et al. (2013) were made with the Wisconsin $\text{H}\alpha$ Mapper (WHAM) telescope, which has sensitivities of a few hundredths of a Rayleigh (see Haffner et al. 2003 for a complete summary of the telescope and survey technique). WHAM has a 1° beam, which is equivalent to a diameter of nearly 1 kpc at the assumed average distance to the MB of 55 kpc. While the WHAM beam is considerably larger than the final resolution of our radio data, at this size, it is less sensitive to small-scale $\text{H}\alpha$ emission stemming from individual H II regions and is optimized to detect faint emission from diffuse ionized gas. For simplicity, we assume an electron temperature of $T_e = 10^4 \text{ K}$ (denoted T_4), as assumed in Barger et al. (2013). Fig. 8 shows the MB region with white contours indicating levels of uncorrected $\text{H}\alpha$ emission from Barger et al. (2013), tracing the 0.06, 0.15 and 1.0 R intensity levels.

Observed $\text{H}\alpha$ intensities are reduced from their intrinsic values due to dust contained within the MB itself and in the MW. These are known as internal and foreground extinction, respectively. We have corrected for both sources of extinction according to table 2 from Barger et al. (2013). We assume that the ‘Join’ and ‘West’ regions have similar interstellar- and local dust content – and therefore an identical total-extinction correction of 28 per cent has been applied. $\text{H}\alpha$ -intensity correction of 22 per cent has been applied to all ‘Wing’ points. For all future analysis and discussion, $\text{H}\alpha$ intensities have been extinction corrected, unless stated otherwise.

We cross-reference the position of each background polarized source with the WHAM data and accept the pointing with the smallest angular separation from our target as the representative $\text{H}\alpha$ brightness for that particular sightline. Because the WHAM survey of the MB is Nyquist sampled, the maximum angular separation allowed is less than 30 arcmin, which corresponds to $\leq 500 \text{ pc}$ at our assumed distance to the MB of 55 kpc. EMs are then derived towards each matched sightline. Mean EMs for each region are listed in Table 4.

5.2 Distribution of ionized medium

In order to estimate the magnetic-field strength along the line of sight through the MB, we assume that there is no correlation between electron density and magnetic-field strength. This has been shown to be a reasonable approximation for typical gas densities associated with the diffuse interstellar medium (Crutcher, Heiles & Troland 2003). Rearranging equation (3), it can be shown that the equation for magnetic field along the line of sight becomes

$$B_{\parallel} = \frac{\phi_{\text{MB}}}{0.812 \bar{n}_e L_{\text{H II}}}, \quad (10)$$

where ϕ_{MB} is the MW-foreground corrected Faraday depth and \bar{n}_e is the mean electron density along the total pathlength of ionized material ($L_{\text{H II}}$).

¹ $1 R = (10^6/4\pi) \text{ photons cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1}$ that is equivalent to $5.7 \times 10^{-18} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ arcsec}^{-2}$ for $\text{H}\alpha$.

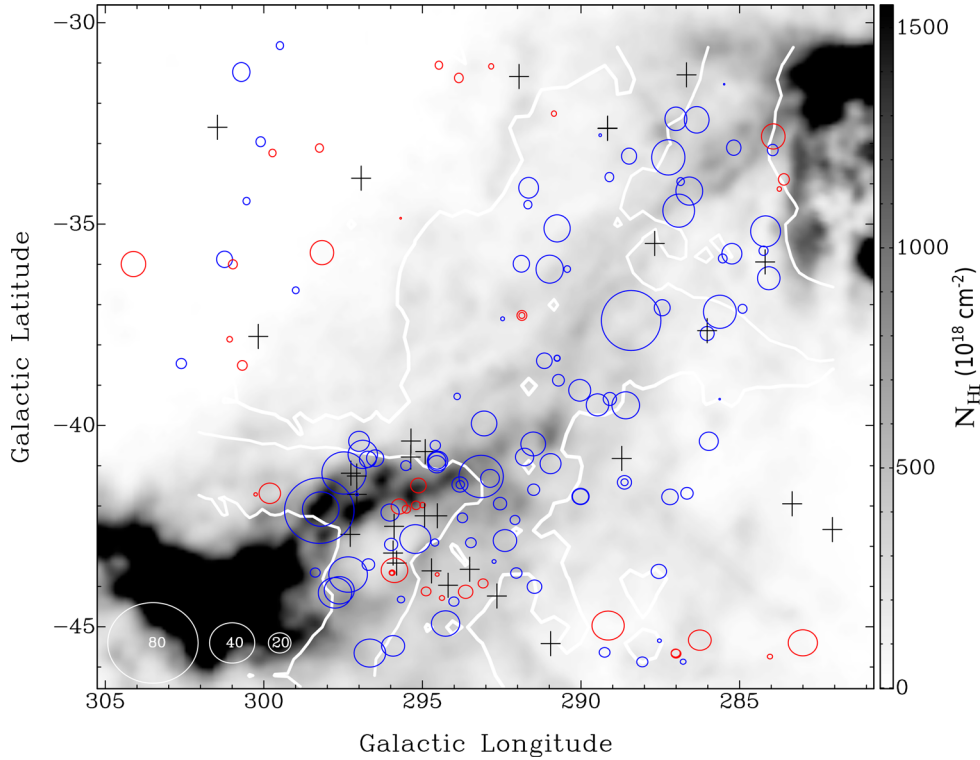


Figure 8. Neutral hydrogen intensity from GASS (Kalberla et al. 2010) in the velocity range of $+100 \leq v_{\text{LSR}} \leq +300$ km s $^{-1}$ overlaid with white contours representing non-extinction corrected H α intensities of 0.06, 0.15 and 1.0 R as measured from WHAM (Barger et al. 2013). Circles represent the foreground-corrected Faraday depth (ϕ_{MB}) values towards each polarized background source. Red and blue circles represent a line-of-sight magnetic field oriented towards and away from the observer, respectively. Black crosses mark the existence of ϕ values that are consistent with zero to $2 \times d\phi$.

Table 4. Table of derived values for polarized sources in all regions of the MB. Column (1) specifies the region of interest, column (2) gives the average H I column density for the region as measured from GASS (Kalberla et al. 2010) and column (3) gives the average extinction-corrected EM from the WHAM data set (Barger et al. 2013). The mean foreground-corrected Faraday depth is given in column (4). Column (5) gives the standard deviation of the foreground-corrected Faraday depth of the region about the mean. The average ionization fraction for each region, assuming the neutral and ionized material is well mixed, as determined by Barger et al. (2013), is listed in column (6). Columns (7–9) give the median coherent magnetic-field strength along the line of sight for each of the ionization geometries discussed in Section 5.2. The errors listed represent the deviation from the 25th and 75th percentiles. The implied random magnetic-field strength, as calculated from equation (19) is given in column (10).

(1) Region	(2) $\langle N_{\text{HI}} \rangle$ ($\times 10^{20}$ cm $^{-2}$)	(3) $\langle \text{EM} \rangle^a$ (pc cm $^{-6}$)	(4) $\bar{\phi}_{\text{MB}}$ (rad m $^{-2}$)	(5) $\sigma(\phi)$ (rad m $^{-2}$)	(6) X (per cent)	(7) $B_{\parallel,1}^*$ ($\mu\text{G } L_5^{-1/2}$)	(8) $B_{\parallel,2}^*$ (μG)	(9) $B_{\parallel,3}^*$ ($\mu\text{G } L_5^{-1}$)	(10) B_r (μG)
Join	3.0	0.283	-13.5	9.3	46	$-0.61^{+0.22}_{-0.41}$	$-0.43^{+0.19}_{-0.15}$	$-0.56^{+0.28}_{-0.57}$	0.11
West	2.8	0.603	-12.3	15.4	46	$-0.33^{+0.29}_{-0.44}$	$-0.27^{+0.23}_{-0.41}$	$-0.23^{+0.20}_{-0.25}$	0.84
Wing	5.0	0.823	-8.7	15.4	29	$-0.26^{+0.32}_{-0.45}$	$-0.29^{+0.35}_{-0.43}$	$-0.27^{+0.21}_{-0.57}$	1.0
-H I Wing	3.6	0.295	-8.1	12.5	24	$-0.39^{+0.44}_{-0.40}$	$-0.40^{+0.46}_{-0.35}$	$-0.44^{+0.48}_{-0.50}$	
-H α Wing	6.8	1.69	-9.7	19.1	21	$-0.06^{+0.14}_{-0.32}$	$-0.07^{+0.16}_{-0.38}$	$-0.08^{+0.17}_{-0.36}$	
Total						$-0.34^{+0.33}_{-0.45}$	$-0.32^{+0.31}_{-0.36}$	$-0.28^{+0.27}_{-0.46}$	

^aThe mean EM is not used in the derivation of magnetic-field strengths. It is listed to give the reader an appreciation of the characteristics of the region.

Often, pulsar dispersion measures ($\text{DM} = \bar{n}_e L_{\text{H I}}$) can be used to construct well-formed estimates of the pathlength and electron density through the different regions. Unfortunately, there are no known pulsars in the MB and very little is known about the morphology and line-of-sight depth of the MB.

Subramanian & Subramanian (2009) argue that the SMC is nearly edge-on, indicating a pathlength through the galaxy of ≥ 5 kpc. If the bulk of the material in the MB had its origins in the SMC, one might expect the depth of the MB to be equally large. Muller et al. (2004) argue that there are numerous obser-

vations throughout the MB that hint at a large line-of-sight depth and Gardiner, Sawa & Fujimoto (1994) estimate the pathlength through regions of the MB $5 \text{ kpc} \lesssim L \lesssim 10 \text{ kpc}$. For simplicity, we parametrize and evaluate the depth of the MB as $L_5 = 5 \text{ kpc}$ and consider the implications of different pathlengths through this parameter, with $1 \lesssim L_5 \lesssim 2$.

Several independent assumptions corresponding to the distribution and geometry of ionized and neutral gas can be made in order to validate our B_{\parallel} measurements. Below, we describe three separate ionized gas distributions and discuss how each might affect derived

magnetic-field strengths. In our discussion, all ionized parameters will be denoted with subscript H II and all neutral gas parameters will be denoted with subscript H I, unless otherwise noted.

5.2.1 Case 1: constant dispersion measure

When estimating the line-of-sight magnetic-field strength, the simplest model of the distribution of material in the MB is one in which the neutral and ionized gas are well mixed. In such a scenario, the bulk of the neutral gas would be distributed across the MB in small clumps, with the ionized medium distributed uniformly amongst the neutral clouds. Therefore, the effective depth of the ionized medium can be expressed as a fraction of the depth of the neutral material, $L_{H II} = f L_{H I}$, where $L_{H I}$ is the depth of the neutral gas and f is the filling factor of ionized gas along the total line of sight (Reynolds 1991).

Little is known of the effective filling factor of ionized gas along the line of sight, but a filling factor of $f=1$ is highly unlikely. Previous work on nearby HVCs in the Leading Arm (McClure-Griffiths et al. 2010) has assumed a filling factor of $f \sim 0.5$ to describe the distribution of the ionized gas and we assume the same value for our analysis. In Section 5.3, we briefly explore the implications of a range of filling factors. Combining the derived EM with our line-of-sight estimates, the DM becomes

$$DM = (EM f L_{H I})^{1/2}. \quad (11)$$

Incorporating the above expression for DM with equation (10), estimates of the magnetic field along the line of sight can be evaluated as

$$B_{\parallel} = \frac{\phi_{MB}}{0.812 (EM f L_{H I})^{1/2}}. \quad (12)$$

This assumption of the geometry of the ionized material in the MB is likely an oversimplification of the actual distribution, which is expected to vary as a function of position along the MB.

5.2.2 Case 2: constant regional ionization fraction

In contrast to Case 1, where we estimated the effective pathlength of the ionized material, here we estimate the free-electron content of a sightline using the ionization fractions (X) across the MB. In order to motivate this approach, a few assumptions must be made. First, we assume that the bulk of the MB material is in the velocity range $+100 \leq v_{LSR} \leq +300$ km s⁻¹ relative to the Galactic Centre (Muller et al. 2003; Putman et al. 2003).

Following from the previous assumption, we also assume that the observed H I depth from GASS (Kalberla et al. 2010) in our selected velocity range probes the entire line-of-sight depth of the MB such that the ionization fraction of a region represents the sum of ionized material in the MB along a given sightline. Previous investigations into the MB have shown this assumption to be reasonable in the diffuse regions of the MB, where observations have shown there is little dust content (Smoker et al. 2000; Lehner et al. 2008). However, there have been observations of molecules in the ‘Wing’ region (Muller et al. 2004; Mizuno et al. 2006; Lehner et al. 2008) and this assumption will serve as a lower limit to our estimates of neutral and ionized gas densities in this region. This second assumption indirectly implies that the neutral and ionized gas are well mixed (i.e. $f \sim 1$; $L_{H I} \simeq L_{H II}$) since any reported ionization fraction is reflective of the pathlength of neutral gas.

Following from these assumptions, the electron density is calculated simply as the ionization fraction multiplied by the neutral gas density

$$\bar{n}_e = \frac{X \langle N_{H I} \rangle}{f L_{H I}}. \quad (13)$$

As with Case 1, the above expression has the underlying premise of $L_{H II} = f L_{H I}$. It follows then that the DM can be written as

$$DM = \frac{X \langle N_{H I} \rangle}{f L_{H I}} L_{H II} = 3.09 \times 10^{18} X \langle N_{H I} \rangle, \quad (14)$$

where the constant 3.09×10^{18} is the conversion factor of pc to cm.

With an expression for the DM, it is now possible to estimate the magnetic field along the line of sight by combining equations (10) and (14),

$$B_{\parallel} = 3.80 \times 10^{18} \left(\frac{\phi_{MB}}{X \langle N_{H I} \rangle} \right). \quad (15)$$

The MB is highly ionized with ionization fractions dependent upon location within the MB (Lehner et al. 2008; Barger et al. 2013). Barger et al. (2013) determined the minimum multiphase ionization fraction across the MB and argued that in the case where the neutral and ionized gas is well mixed, the average ionization fraction in the region of the diffuse MB is $X \simeq 46$ per cent, and $X \simeq 29$ per cent in the ‘Wing’ region. As these ionization fractions represent the average values calculated over the entire region, the values do not represent small-scale variations in the distribution of material. In contrast, the spectroscopic work of Lehner et al. (2008) found ionization fractions as high as $X \simeq 90$ per cent along three sightlines corresponding to the ‘Join’ and ‘West’ regions. As their sightlines probed localized distributions, this approach would have been susceptible to small-scale enhancements.

Motivated by the wide variability in ionization fractions, we choose to evaluate the ionization level of the various regions individually. In the region of the Wing, we compare the H I and H α column densities from table 3 in Barger et al. (2013), to calculate a multiphase ionization fraction of $X \simeq 29$ per cent in the ‘Wing’. We evaluate the ‘Join’ and ‘West’ regions at an ionization fraction of $X \simeq 46$ per cent, assuming that the ‘Join’ and ‘West’ regions host similar distributions of material. Evaluation of this ionization level makes it simple to explore the range of possible magnetic-field strengths.

In the region of the SMC-Wing, there is a clear variation of H I column densities as well as H α intensities (Fig. 8). Following Barger et al. (2013), we choose to break up the Wing into two regions corresponding to the relative H α brightness. If a sightline is associated with an uncorrected H α brightness larger than 0.15 R, this region is classified as the ‘H α -Wing’ and assigned an H I column density of 6.8×10^{20} cm⁻², else we consider the region to be the ‘H I-Wing’ and evaluate it as having an H I column density of 3.6×10^{20} cm⁻². Furthermore, assuming well-mixed neutral and ionized gas populations (e.g. $L_{H I} \simeq L_{H II}$), we use the mass estimates of Barger et al. (2013) to calculate an ionization fraction of $X \simeq 24$ and $X \simeq 21$ for the ‘H I-’ and ‘H α -Wing’, respectively. A summary of region parameters is given in Table 4.

This assumed geometry of the distribution of ionized gas is similar to Case 1 (Section 5.2.1), in that it requires the neutral and ionized media to be well mixed. However, in this model, our greatest approximation is the mean ionization fraction for a given region of the MB. Although not stated explicitly in equation (15), this B_{\parallel} estimate does have a dependence on the assumed pathlength of ionized material through the derivative of the total ionized mass of

the region and subsequently implied X , which is outlined in Barger et al. (2013).

5.2.3 Case 3: ionized skin

Ionizing photons that have escaped the MW and the Magellanic Clouds have the potential to ionize the outer layers of the MB (Fox et al. 2005; Barger et al. 2013). In this possibility, the distribution of the thermal electrons is that of an ionized skin, rather than mixed with the neutral gas, as we assumed in Cases 1 and 2. In order to explore this third scenario, we assume that the neutral hydrogen is girt by a fully ionized skin at the same temperature and pressure, the density of which will be $\bar{n}_e = \bar{n}_{\text{H}^+}/2$ (Hill, Haffner & Reynolds 2009). This condition requires that the neutral and ionized media have had enough time to come into pressure equilibrium, which we assume for our analysis.

In the ionized skin, the line-of-sight depth can be derived from our density assumption combined with $\text{H}\alpha$ brightnesses:

$$L_{\text{H}\alpha} = \text{EM} n_e^{-2} = \text{EM} \left(\frac{\bar{n}_{\text{H}^+}}{2} \right)^{-2} = 4 \text{EM} \left(\frac{f L_{\text{H}\alpha}}{\langle N_{\text{H}^+} \rangle} \right)^2, \quad (16)$$

where the discussion for the evaluation of n_e is given in the previous model (equation 13). We can now combine equation (16) with our density estimates (equation 13) to find an expression for DM:

$$\text{DM} = \frac{2 \text{EM} f^2 L_{\text{H}\alpha}}{\langle N_{\text{H}^+} \rangle}. \quad (17)$$

Substituting this expression for DM into equation (10), the equation for the magnetic-field strength along the line of sight in an ionized skin becomes

$$B_{\parallel} = \frac{\phi_{\text{MB}} \langle N_{\text{H}^+} \rangle}{1.6 \text{EM} f^2 L_{\text{H}\alpha}} \quad (18)$$

where the line of sight of the neutral medium is in units of cm.

In the case of an ionized skin, the pathlength of the ionized medium is expressed explicitly in terms of our two assumptions: first, that the neutral and ionized media are in pressure equilibrium and secondly, that the filling factor of the neutral medium is $f_{\text{H}^+} \simeq 1$ along the effective depth of the MB. Therefore, we argue that for the above thin-skin approximation, $f \simeq 1$.

5.3 Summary and comparison of ionization cases

We evaluate each of the aforementioned cases for a line-of-sight pathlength of $L_{\text{H}^+} = L_5 = 5$ kpc. As shown in Table 4 and Fig. 9, each of the cases results in similar estimates in line-of-sight magnetic-field strengths for the entire MB, with median values all near $B_{\parallel} \simeq 0.3 \mu\text{G}$. By comparison, individual regions show a larger scatter between derived magnetic-field strengths.

Our sample of Faraday depths is skewed towards the negative (as seen in Fig. 8); therefore, it follows that the derived field strengths are distributed in kind. By completing a skewness test, we find that the B_{\parallel} distribution resulting from Case 1 is skewed towards negative values with a 3.2σ confidence level. It follows that Case 2 is skewed negative to 1.7σ and Case 3 to 4.2σ significance. This skew can be seen most clearly in Fig. 9. Due to this skew towards negative values, the B_{\parallel} values quoted in Table 4 represent the median magnetic-field strengths, where the median statistic is more robust against outliers. Along with the median value, we list the deviation from the first and third quartile (Q_1 and Q_3 , respectively), which represents the 25th and 75th percentile values in the distribution. The derived magnetic-

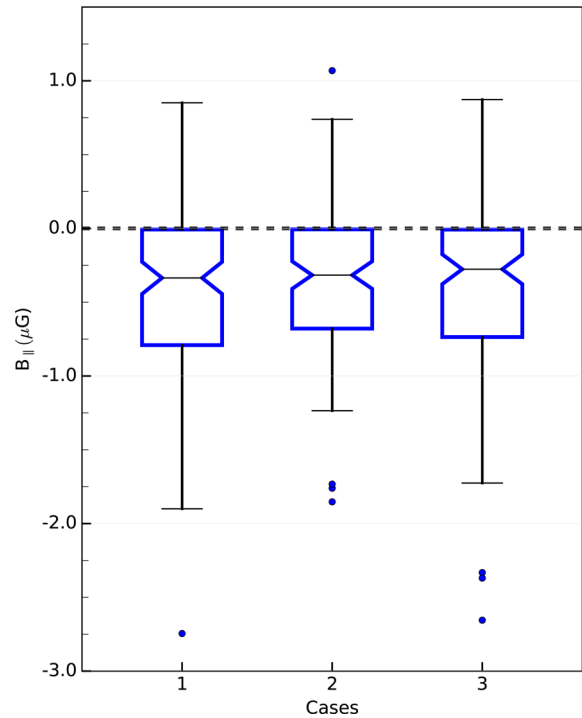


Figure 9. Box-and-whisker plots of coherent line-of-sight magnetic field measurements towards all sightlines through the MB for all cases listed in Section 5.2. The height of each box marks the IQR of each distribution while notches mark the median position. The caps at the end of the whiskers represent the 5th and 95th percentiles. The outliers in each population are shown as dots above and below the whiskers. There is a dashed line at $B_{\parallel} = 0$ to clarify the distribution of positive and negative magnetic field orientations.

field strengths are best summarized by Fig. 9, where the bound region denotes the interquartile range (IQR), defined as $\text{IQR} = Q_3 - Q_1$.

As noted in our discussion of the ionization models, the largest uncertainty in our B_{\parallel} measurements comes from the unknown geometry of the MB along the line of sight, namely the uncertainty in L_{H^+} , X , and f . With that in mind, we aim to compare all models by their dependence on our depth assumptions and use of measured quantities.

Cases 1 and 2 are built from the same oversimplified picture of well-mixed neutral and ionized gas distributions along the line of sight. Case 1 uses only the measured EM with a largely unconstrained filling factor, f . Exploring a range of f for Case 1 shows that a 20 per cent change in f (i.e. $0.3 \leq f \leq 0.7$) results in less than a $0.1 \mu\text{G}$ change in the median line-of-sight magnetic-field strength. In contrast, Case 2 takes advantage of more information, using both calculated ionization fractions and measured $\langle N_{\text{H}^+} \rangle$. Contrasting these first two ionized gas distributions, Case 3 has the ionized material distributed as an ionized skin. This geometry requires that the neutral and ionized gas to be in pressure equilibrium in order to be physical. It is possible that this condition could be met in the ‘Join’ region; however, it is likely that this is inappropriate for regions in the ‘Wing’ due to energy being injected from ongoing star formation (e.g. Noël et al. 2015).

We show that Case 2 has least dependence on an assumed pathlength through the MB and filling factor. As we mentioned in Section 5.2.2, the actual ionization fractions across the MB may be higher than our evaluated estimates. Increasing the ionization fraction by a factor of 2 implies that the line-of-sight

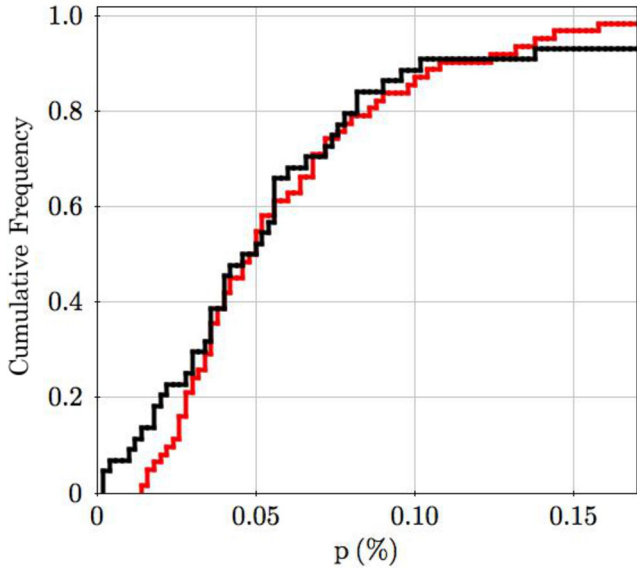


Figure 10. Cumulative histogram comparing observed polarization fractions (p) for sources on (red) and off the Bridge (black). Sources associated with the ‘Wing’ region are not included in this distribution due to the source selection bias towards sources with high polarization. This figure has been truncated at $p = 17$ per cent for clarity.

magnetic-field strength is half the current value – i.e. $B_{\parallel}(X = 90 \text{ per cent}) \simeq -0.17 \mu\text{G}$. The following discussion will be carried out using the magnetic field estimates derived from Case 2 and all parameters reported in Table 4, unless specified otherwise.

6 DISCUSSION

In this section, we consider the implications of the observed Faraday depth values and magnetic-field strengths in the MB and explore possible origins of the coherent magnetic structure.

6.1 The turbulent magnetic field

Ongoing star formation in the MB (e.g. Noël et al. 2015) will make any existent regular magnetic field to become turbulent and random. An increase in random motion would also depolarize any background polarized light, proportional to the level of turbulence. If the magnetic field observed in the MB were sufficiently turbulent, one would expect that the polarization of sources associated with the MB would exhibit higher levels of depolarization, and thus have lower values for the observed fractional polarization.

We explore the consequences of the turbulent field by comparing the observed polarization fraction for populations of sources on and off the MB. Fig. 10 shows a cumulative histogram comparing the observed polarization fractions. We choose not to include sources associated with the ‘Wing’ region due to the source selection bias that favours highly polarized sources. The two source populations show no statistically significant differences in the observed fractional polarization, indicating there is no correlation between source location and turbulence of the foreground magnetic field.

If the turbulence in the field is not strong enough to depolarize the background signal completely, it is still possible to investigate the mean Faraday dispersion (σ_{ϕ}) as fitted by our qu -fitting routine. We compare the values for sources on-Bridge and off-Bridge, under the hypothesis that sources on the MB would exhibit higher σ_{ϕ} if there are more coherent and/or turbulent cells located in the MB

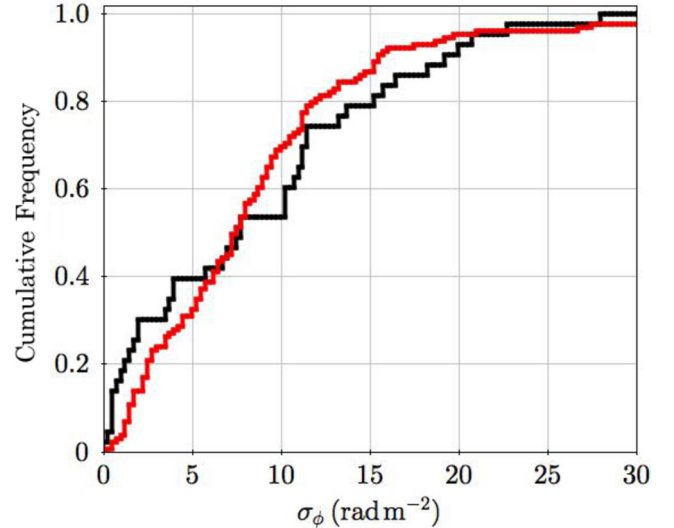


Figure 11. Cumulative histogram of best-fitting Faraday dispersion values (σ_{ϕ}) fitting to a single, simple Faraday-rotating source with foreground depolarization for sources on (red) and off the Bridge (black). The figure has been truncated at $\sigma_{\phi} = 30 \text{ rad m}^{-2}$ for clarity.

when compared to the MW. Fig. 11 shows a cumulative histogram of the best-fitting Faraday dispersion values for all points on and off the MB. We carry out a two-sample Anderson–Darling test with both σ_{ϕ} populations and find that we cannot reject the null-hypothesis of the two samples being drawn from the same distribution and conclude that any turbulence in the MB magnetic field cannot be differentiated from that in the MW.

In the MW, LMC and SMC, it has been shown that the random component of the magnetic field dominates the total field strength (Beck 2000; Gaensler et al. 2005; Mao et al. 2008). To estimate the random magnetic-field strength, we choose a similar approach to Mao et al. (2008), who assume that the coherent magnetic field does not change as a function of position, but that any change in observed ϕ is due to turbulence. It is then possible to estimate the mean random magnetic-field strength for each observed region as

$$B_r = \frac{3l_0}{L_5} \sqrt{\left(\frac{\sigma(\phi)}{0.812 \bar{n}_e l_0}\right)^2 - \left(\frac{B_{\parallel}^* \Delta L}{l_0}\right)^2}, \quad (19)$$

where $\sigma(\phi)$ is the standard deviation of Faraday depths in the region of interest, \bar{n}_e is the average electron density in units of cm^{-3} and B_{\parallel}^* is the median coherent magnetic-field strength along the line of sight in μG . ΔL is the standard deviation of the pathlength through the ionized medium in pcs and characterizes the uncertainty in the depth of the MB. l_0 is the linear scale of an ‘RM-cell’ in units of pc, such that $n \sim L_5/l_0$, where n is the number of cells in a single line of sight.

As stressed in our derivation of the coherent magnetic-field strength, little is known about the morphology of the MB, leaving the estimations for pathlength to be our largest uncertainty. We estimate the standard deviation of the width of the MB (ΔL) to be 1 kpc. Gaensler et al. (2005) show that RM-cells in the LMC are of order ~ 100 pc, and we adopt a similar value for our analysis. As we have done in Case 2 (Section 5.2.2), we consider \bar{n}_e to be related to the column density of H I as $\bar{n}_e = (X \langle N_{\text{HI}} \rangle) / L_5$ where L_5 is the line-of-sight depth of the MB in cm and X is the ionization fraction of the region.

As discussed in Section 4.1, our correction for the foreground MW Faraday depth does not account for the intrinsic Faraday depth of the source. We minimize any resultant effects by subtracting the scatter of intrinsic extragalactic Faraday depths $\sigma(\phi) = 6 \text{ rad m}^{-2}$ (Schnitzler 2010) from our regional Faraday depth standard deviations.

Using the above estimates and the values listed in Table 4, we derive the implied random magnetic-field strengths of each region, the results of which are summarized in Table 4. We find that the turbulent field dominates the ordered component in the regions of the ‘Wing’ and ‘West’. Intriguingly, this does not hold in the ‘Join’ region. Perhaps this indicates that our pathlength estimates are unrealistic or that our overarching assumptions are unviable. However, the ‘Join’ region is furthest from any ongoing star formation. This fact, combined with our aforementioned turbulence null-hypotheses, suggests that the random field may not dominate the large-scale magnetic field in the diffuse MB.

6.2 Estimating the total magnetic field of the wing

Recent work by Lobo Gomes et al. (2015) mapped coherent field lines in the plane-of-the-sky ($B_{c,\perp}$) in the SMC and Wing using optical polarized starlight. They argued that there exists a significant fraction of sightlines that exhibit a magnetic field that points in the direction of the MB towards the LMC of order $B_{\perp} = (0.947 \pm 0.079) \mu\text{G}$.

We combine their measurements with our estimations for B_{\parallel} to estimate the total coherent magnetic-field strength ($B_{c,T}$) in the Wing

$$B_{c,T}^2 = B_{c,\perp}^2 + B_{c,\parallel}^2. \quad (20)$$

We find an implied total magnetic-field strength of $B_{c,T} \simeq 1 \mu\text{G}$ in the region of the Wing that implies that the ordered magnetic field in the Wing is dominated by the plane-of-the-sky component. We note that our uncertainty estimates imply a large range of coherent field strengths. For the sake of brevity, we omit the implications of all possible field strengths from our discussion. This field strength is within the range of magnitudes expected if the field were to have originated from the SMC as Mao et al. (2008) estimated total coherent magnetic field in the SMC to be $B_T \simeq 1.7 \pm 0.4 \mu\text{G}$.

6.3 The pan-Magellanic Ffield

The possible existence of a large-scale magnetic field that permeates the entire Magellanic System (the ‘pan-Magellanic field’, pM field) was first introduced by Mathewson & Ford (1970a) and Schmidt (1970). Furthermore, Schmidt (1970) argued that the existence of such a field suggested that the fields observed in the LMC and SMC shared a common origin. Continued investigations into the nature of the magnetism across Magellanic System were carried out (e.g. Schmidt 1976; Mathewson et al. 1979; Wayte 1990) strengthening the case for the existence of the pM field. More recently, Mao et al. (2008) and Lobo Gomes et al. (2015) note the potential alignment of the SMC magnetic field with the MB, and Mao et al. (2012) argue the same for the LMC. However, all previous research has been confined to high-density regions in the LMC and SMC.

If the pM field exists, it is expected to be dominated by the plane-of-the-sky component, just as the fields associated with the LMC and SMC have been observed to be (Mao et al. 2012; Lobo Gomes et al. 2015). However, the observation of negative Faraday depths across the MB implies a non-trivial line-of-sight component. We argue that this directional component was anticipated, as the SMC

is located further away from the MW than the LMC (~ 60 and ~ 50 kpc, respectively; Walker 1999). Therefore, our observation of a Faraday-depth signal spanning the entirety of the MB may be the first direct evidence of the pan-Magellanic field.

Before we can confirm the existence of the pan-Magellanic field, it is important to understand if the coherent fields associated with the SMC and LMC can account for the observed Faraday depth signal seen to span the entire MB. Below we investigate the possible origins of the observed coherent magnetic field and how it might relate to the ‘pan-Magellanic field’ hypothesis. We assume that the observed magnetic field has been frozen in to the tidally stripped gas for this discussion, and address any evidence towards the contrary at the end of this section.

6.3.1 The SMC-Wing field

Mao et al. (2008) estimate that the SMC has a line of sight and total magnetic-field strength of $B_{\parallel} = -0.19 \pm 0.06$ and $B_{Tot,C} = 1.6 \pm 0.4 \mu\text{G}$, respectively. All ionization models discussed in our paper imply a line-of-sight magnetic-field strength that is consistent with this estimate in the ‘Wing’ region, within the IQR (Table 4). We note that our average median value between all Cases is $\overline{B_{\parallel}}^* = -0.27 \mu\text{G}$, a value that is larger than what was observed in the SMC. However, our estimates are again based largely on the assumed geometry of the MB. If the MB is oriented predominantly along the line of sight, as is the orientation of the SMC, then the ionized pathlength would be larger than L_5 . Doubling the line-of-sight depth through the Bridge ($2L_5$) results in a decreased estimated magnetic-field strength of $\overline{B_{\parallel}}^* = -0.20 \mu\text{G}$, a value that agrees well with the magnetic field estimates of Mao et al. (2008).

If the SMC magnetic field is responsible for the observed MB field, it is also possible to estimate the expected MB ϕ signal. Using Case 1 (Section 5.2.1) and an average pathlength of $\overline{L} = L_5$, a filling factor of $f=0.5$ and estimates of the average electron density from the EMs from Table 4, a coherent line-of-sight magnetic field of B_{\parallel} would manifest itself as a Faraday depth of $\phi \sim -7$ and -4 rad m^{-2} in the ‘Wing’ and ‘Join’ regions, respectively. This approximation is roughly consistent with the observed median ϕ in the Wing, but appears to contradict the mean Faraday depth measured in the ‘Join’ region, which returns an average Faraday depth that has a magnitude more than three times the expected value ($\phi_{MB} = 12.5 \text{ rad m}^{-2}$). To achieve the observed mean Faraday depth in the ‘Join’ region with the SMC $\overline{B_{\parallel}}$ requires an effective pathlength of ~ 27 kpc, which is unlikely to be physical.

This contradiction can be accounted for if the orientation of the magnetic field changes as a function of position along the MB. If in the ‘Join’ region, the coherent field has been rotated such that a larger fraction of the total field lies along the line of sight, one would observe a larger ϕ_{MB} and derive a stronger B_{\parallel} . Indeed, this is what is observed in the ‘Join’ region.

We argue that the magnetic field present in the ‘Wing’ region is consistent with a field that was created in the SMC and pulled into the MB. It is possible that the inherited field stretches as far as the ‘Join’ region. However, it requires that the turbulent field to be less significant than that associated with the ‘Wing’ region. It is then plausible that the pan-Magellanic field is governed by the geometry of the coherent field in the SMC. We do not extend this analysis to the ‘West’ region as the field associated with this region is possibly an extension from the LMC rather than the SMC. We now explore this possibility.

6.3.2 The LMC-West field

Fig. 8 shows a nearly consistent, negative ϕ -value in the region nearest the LMC. Leading to this region, Gaensler et al. (2005) show the bulk of the polarized sources in the nearest portion of the LMC also have negative RM-values² after MW-foreground correction. Contrasting this, the majority of the RMs associated with the LMC have positive values.

In the plane-of-the-sky magnetic field, both Wayte (1990) and Mao et al. (2012) note that there are regions of the LMC where the magnetic field appears to align with the direction towards the SMC.

If the MB-field is built from the magnetized material that originated in both Magellanic Clouds, the LMC contribution is likely associated with the tidal filaments ($l \simeq 285.7$, $b \simeq -33.7$), first identified by Haynes et al. (1991). Mao et al. (2012) estimate that the tidal filaments stemming from the LMC have a magnetic-field strength of $B_T = 11 \mu\text{G}$. They argue that the similar magnitude of off-source and on-source Faraday depths signifies that the line-of-sight magnetic field in this region is negligible.

We measure an average $\bar{\phi}_{MB} = -21 \text{ rad m}^{-2}$ in the region of the filament discussed above. Although we have fewer off-source points, we find a scatter of the nearest 10 sources to be $\sigma \sim 2 \text{ rad m}^{-2}$. For there to be no line-of-sight magnetic field in this region, the average electron density would have to be zero, which is unphysical. We estimate the B_{\parallel} of the tidal filament using the same estimated pathlength as Mao et al. (2012), $L_{H_{II}} = 800 \text{ pc}$ and Case 1 (Section 5.2). The implied line-of-sight magnetic field in this region is $B_{\parallel} = 1.2 \mu\text{G}$, a value that is twice as strong as our initial estimates ($L_{H_{II}} = fL_5$) for the same region.

This exercise demonstrates two things: first, it is possible that the Faraday depth values we see in the region nearest the LMC could be a consequence of a coherent magnetic field having been stripped from the LMC. Secondly, with different pathlength estimates, the magnetic field contribution from the LMC could be much higher than our initial estimates, implying a stronger total magnetic-field strength in the MB.

6.3.3 Does the pan-Magellanic field exist?

We have discussed the implications of the known magnetic fields of the SMC and LMC as they pertain to our observed ϕ_{MB} and $B_{\parallel,MB}$ in an attempt to justify the assumption that the observed MB magnetic field originated from both galaxies. We posit that the dominant magnetic field component is in the plane of the sky, a claim that is consistent with what has been previously argued in all discussions of the pan-Magellanic field.

However, the previous discussions and implications of the SMC and LMC magnetic fields assumed that the observed MB field is a combination of magnetic fields that have been drawn out of the LMC and SMC. Below, we briefly explore if the observed coherent field in the MB could have been formed in situ.

The $\alpha - \omega$ dynamo – which is believed to be the mechanism responsible for the observation of coherent magnetic fields on the scales of galaxies – requires too large a time-scale to explain the existence of a coherent field in the young MB. By comparison, the cosmic ray driven dynamo works on much shorter time-scales. However, both of these dynamos require there to be differential rotation in the MB, which has not been observed. Therefore, these

mechanisms cannot be responsible for the magnetic field in the MB. By contrast, the typical amplification time of the fluctuating dynamo is $10^6 - 10^7 \text{ yr}$, a time-scale that is favourable given the age of the MB. However, this mechanism creates turbulent or incoherent magnetic fields and cannot be responsible for the observed ordered field in the MB.

While we do not address the origins of the magnetic fields in the LMC and SMC, the standard magnetic-field creation mechanisms cannot explain the existence of a coherent magnetic field in the young tidal remnant. We therefore conclude that the magnetic field in the MB is a consequence of field lines having been dragged out of the LMC and SMC with an overarching field geometry that has been determined by the orientation of the parent galaxies. This shared magnetic history links the two Magellanic Clouds and establishes the existence of the pan-Magellanic field.

All previous detections of tidal bridges with corresponding polarization have been detected through polarized continuum emission emanating from tidal regions (e.g. Condon et al. 1993; Nikiel-Wroczyński et al. 2013a,b). Our detection of a coherent magnetic field in the MB was made using Faraday rotation observations through a non-continuum foreground, and is the first ever such detection for any tidal bridge. This may imply that magnetic fields are an early influence on the evolution of galaxy interactions. If magnetic fields do affect early galaxy interactions, the existence of magnetic fields in tidal remnants could explain the observation of coherent magnetic fields in tidal dwarfs (Nikiel-Wroczyński et al. 2013a) and would suggest the existence of magnetic fields in more diffuse tidal features, such as the Magellanic Stream and Leading Arm.

7 SUMMARY

We have presented Faraday rotation data for 167 extragalactic polarized sources and observe a coherent magnetic field towards the MB. Each source in our catalogue has well-determined polarization ($\mathcal{P} \geq 8\sigma_{\mathcal{P}}$). Using an MCMC approach to fitting observed complex polarization spectra, we were able to recover the polarization parameters of each source to high confidence.

We have demonstrated that the observed Faraday depths of sources ‘on’ and ‘off’ the MB are inherently different and have attributed this disparity to the existence of a large-scale, coherent magnetic field within the MB. We assumed a line-of-sight depth through the MB of 5 kpc and explored different distributions of ionized gas. The median line-of-sight magnetic fields derived from these approximations are all consistent with $B_{\parallel} \simeq 0.3 \mu\text{G}$, where the uniform field is directed away from us. We stress that little is known about the distribution of ionized gas within the MB and the implied magnetic field is dependent upon this constraint.

The MB is a tidal remnant that we argued has no known means for creating a coherent field on the scales observed. Therefore, we concluded that the magnetic fields of the LMC and SMC have been tidally stripped along with the neutral gas emanating from these galaxies to form the MB. The implied line-of-sight magnetic-field strength in the MB region nearest the SMC, which is where the majority of the gas of the MB is believed to have originated, is consistent with observed line-of-sight component of this galaxy. We have argued that the magnetic field associated with the LMC and its polarized filaments has also been pulled into the MB and is likely responsible for the observed Faraday rotation in the region nearest these features.

This work represents the first observational confirmation of the pan-Magellanic field – a coherent magnetic field spanning the

² Previous studies explicitly use the term RM, rather than Faraday depth, to express the magnitude of observed Faraday rotation.

entirety of the MB with a history and evolutionary fate that is tied to that of the Magellanic System.

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APPENDIX A: TABLE OF SYMBOLS

Table A1. List of symbols used in this paper and their meaning.

Symbol	Physical quantity
B_{\parallel}	Measured magnetic-field strength along the line of sight in units of μG
B_{\parallel}^*	Median magnetic-field strength along the line of sight in units of μG
$B_{c,T}$	Total coherent magnetic-field strength, in units of μG
B_r	Random magnetic-field strength, in units of μG
DM	Measured dispersion measure for a specific sightline in units of pc cm^{-3}
$\langle \text{EM} \rangle$	Average emission measure for specified region, in units of pc cm^{-6}
EM	Measured emission measure along a specific sightline, pc cm^{-6}
f	Volume filling factor of gas such that the effective pathlength of gas with a characteristic density n_0 is $f \times n_0$
X	Ionization fraction
I, Q, U, v	Observed Stokes parameters, with units of mJy
$I_{\text{H}\alpha}$	Intensity of $\text{H}\alpha$ emission, with units of Rayleighs
L_5	5000 pc. The nominal line-of-sight depth of the MB.
$L_{\text{H I}}$	Estimated line-of-sight depth of neutral hydrogen, in units of pc
$L_{\text{H II}}$	Estimated line-of-sight depth of ionized material, in units of pc
ΔL	Estimated standard deviation in line-of-sight depth of the MB, in units of pc
IQR	Interquartile range, defined as the range of values between the 25th and 75th percentiles
l_0	Typical cell size along line of sight, in units of pc
λ^2	The square of the observed wavelength, in units of m^2
$\langle N_{\text{H I}} \rangle$	Average H I column density, in units of cm^{-2}
$\bar{n}_{\text{H I}}$	Average neutral gas density, calculated as $\langle N_{\text{H I}} \rangle / L_{\text{H I}}$ in units of cm^{-3}
\bar{n}_e	Average free-electron density, in units of cm^{-3}
\mathcal{P}	Polarized intensity in units of mJy/beam
p	Observed polarized fraction
p_0	Intrinsic polarized fraction
ϕ_{corr}	Faraday depth for which the foreground, MW contribution has been subtracted, in units of rad m^{-2}
ϕ_{MB}	Faraday depth of the MB, in units of rad m^{-2}
ϕ_{obs}	Observed Faraday depth in units of rad m^{-2}
$\Delta\phi$	Uncertainty estimate in ϕ from fitting algorithm, qu -fitting, in units of rad m^{-2}
$\bar{\phi}$	Mean Faraday depth in units of rad m^{-2}
σ_{ϕ}^2	Variance of Faraday depths on scales smaller than the synthesized beam, in units of rad m^{-2}
$\sigma(\phi)$	Standard deviation of an ensemble of Faraday depth values for a specified region, in units of rad m^{-2}
Ψ	Observed polarization angle, defined as $0.5 \arctan \frac{U}{Q}$
Ψ_0	Intrinsic polarization angle at the source of emission
$\mathcal{Q}_1, \mathcal{Q}_2$	First and third quartile, defined as the 25th and 75th population percentile value
RM	The classical rotation measure, defined as $(\Psi_0 - \Psi) / \lambda^2$, in units of rad m^{-2}
$\sigma_{\mathcal{P}}$	Rms error in extracted polarized intensity, in units of mJy b^{-1}
q, u	Fractional linear polarized Stokes Q and U parameters, units of per cent
T_4	Assumed temperature of 10^4 K for the ionized medium

APPENDIX B: TABLE OF DERIVED FARADAY DEPTHS

Table B1. Table of the Faraday depth values for each polarized source used in our analysis. Sources preceded by an ‘*’ indicate targets that are considered to be off the Bridge, whereas those without an asterisk are considered to be on-Bridge sources. Columns 1 and 2 give the position of the source in galactic longitude and latitude. Columns 3–6 lists the best-fitting values returned from the $q - u$ fitting routine, namely the intrinsic polarization fraction, intrinsic polarization angle, observed Faraday depth and Faraday dispersion. Each of the uncertainties represents the 1σ standard deviation in parameter space. Column 7 lists the corresponding Faraday depth of each source once the Faraday rotation due to the MW foreground has been corrected for. The uncertainty calculation for this value is described in detail Section 4.1.

l ($^{\circ}$)	b ($^{\circ}$)	p_0 (per cent)	Ψ_0 ($^{\circ}$)	ϕ_{raw} (rad m^{-2})	σ_{ϕ} (rad m^{-2})	ϕ_{corr} (rad m^{-2})
* 282.073	-42.586	0.025 \pm 0.001	10.5 \pm 2.7	24.9 \pm 1.5	0.9 \pm 0.8	-1.2 \pm 1.5
* 283.003	-45.398	0.068 \pm 0.001	3.1 \pm 0.7	47.9 \pm 0.5	11.0 \pm 0.5	25.9 \pm 0.5
* 283.340	-41.945	0.043 \pm 0.001	44.0 \pm 1.9	24.2 \pm 1.1	1.0 \pm 0.9	-2.0 \pm 1.1
283.601	-33.891	0.078 \pm 0.004	0.9 \pm 2.7	47.8 \pm 1.5	10.0 \pm 1.3	11.4 \pm 1.6
283.747	-34.130	0.097 \pm 0.005	94.9 \pm 2.1	40.6 \pm 1.2	7.5 \pm 1.5	4.6 \pm 1.2
283.937	-32.823	0.045 \pm 0.001	3.8 \pm 0.6	62.7 \pm 0.5	9.0 \pm 0.4	25.0 \pm 0.6
283.953	-33.160	0.060 \pm 0.002	16.5 \pm 1.9	25.9 \pm 1.1	1.1 \pm 1.0	-11.3 \pm 1.1
* 284.041	-45.740	0.044 \pm 0.001	6.2 \pm 1.0	25.7 \pm 0.7	16.6 \pm 0.4	4.6 \pm 0.7
284.078	-36.344	0.073 \pm 0.004	19.9 \pm 2.5	10.2 \pm 1.5	11.6 \pm 1.1	-22.9 \pm 1.5
284.178	-35.180	0.072 \pm 0.002	14.7 \pm 1.6	3.8 \pm 0.9	1.3 \pm 1.1	-30.7 \pm 0.9
284.193	-35.940	0.032 \pm 0.002	93.9 \pm 2.2	32.5 \pm 1.6	9.7 \pm 1.5	-1.0 \pm 1.6
284.245	-35.661	0.162 \pm 0.015	118.6 \pm 4.0	24.6 \pm 2.3	13.3 \pm 1.6	-9.3 \pm 2.4
284.904	-37.103	0.070 \pm 0.008	121.7 \pm 5.4	22.8 \pm 3.2	14.4 \pm 1.8	-8.9 \pm 3.2
285.180	-33.102	0.053 \pm 0.004	43.2 \pm 2.9	21.4 \pm 1.9	14.5 \pm 1.2	-15.2 \pm 1.9
285.244	-35.736	0.060 \pm 0.003	16.8 \pm 1.9	12.8 \pm 1.0	5.9 \pm 1.6	-20.4 \pm 1.1
285.485	-31.527	0.334 \pm 0.009	44.5 \pm 1.2	37.0 \pm 0.7	7.5 \pm 0.7	-1.5 \pm 0.7
285.532	-35.854	0.133 \pm 0.006	42.2 \pm 2.0	24.0 \pm 1.1	8.0 \pm 1.1	-8.9 \pm 1.1
285.621	-37.178	0.132 \pm 0.007	122.4 \pm 2.5	-1.7 \pm 1.4	11.6 \pm 1.0	-32.9 \pm 1.4
* 285.625	-39.347	0.064 \pm 0.001	143.9 \pm 0.6	26.8 \pm 0.3	13.9 \pm 0.2	-1.6 \pm 0.3
* 285.970	-40.392	0.075 \pm 0.007	87.3 \pm 3.6	8.7 \pm 2.0	7.4 \pm 2.6	-18.2 \pm 2.0
286.019	-37.718	0.049 \pm 0.002	134.3 \pm 1.9	16.3 \pm 1.4	15.6 \pm 0.8	-14.0 \pm 1.4
286.022	-37.647	0.248 \pm 0.011	12.3 \pm 2.9	27.5 \pm 1.4	2.5 \pm 1.9	-2.9 \pm 1.4
* 286.253	-45.332	0.400 \pm 0.035	177.6 \pm 7.5	40.6 \pm 3.4	3.6 \pm 2.5	20.2 \pm 3.4
286.352	-32.410	0.158 \pm 0.018	147.1 \pm 4.2	10.6 \pm 2.0	8.8 \pm 2.4	-26.3 \pm 2.1
286.582	-34.181	0.047 \pm 0.001	154.4 \pm 0.7	6.6 \pm 0.4	1.3 \pm 1.0	-27.9 \pm 0.5
* 286.660	-41.685	0.101 \pm 0.002	118.0 \pm 1.1	13.3 \pm 0.6	1.8 \pm 1.3	-11.6 \pm 0.6
286.672	-31.294	0.291 \pm 0.013	124.5 \pm 2.4	42.8 \pm 2.3	30.6 \pm 0.9	4.6 \pm 2.3
* 286.782	-45.866	0.076 \pm 0.001	109.4 \pm 0.8	14.4 \pm 0.4	3.5 \pm 1.1	-5.1 \pm 0.4
286.858	-33.944	0.082 \pm 0.006	150.3 \pm 3.1	26.9 \pm 1.6	6.9 \pm 2.1	-7.8 \pm 1.6
286.919	-34.667	0.324 \pm 0.024	71.1 \pm 3.2	0.9 \pm 1.7	11.5 \pm 1.3	-32.9 \pm 1.8
* 287.005	-45.668	0.114 \pm 0.001	89.0 \pm 0.4	28.2 \pm 0.2	0.5 \pm 0.4	8.6 \pm 0.2
287.005	-32.386	0.070 \pm 0.002	175.3 \pm 2.1	13.3 \pm 1.1	1.3 \pm 1.1	-23.3 \pm 1.1
* 287.015	-45.653	0.117 \pm 0.001	88.9 \pm 0.4	27.6 \pm 0.2	3.8 \pm 0.5	7.9 \pm 0.2
* 287.195	-41.776	0.070 \pm 0.006	73.3 \pm 3.8	9.0 \pm 2.3	10.9 \pm 2.0	-15.5 \pm 2.3
287.249	-33.344	0.094 \pm 0.003	130.9 \pm 1.3	0.2 \pm 0.7	11.5 \pm 0.5	-35.0 \pm 0.8
287.439	-37.078	0.042 \pm 0.001	160.9 \pm 1.4	14.4 \pm 0.8	9.6 \pm 0.7	-16.0 \pm 0.8
* 287.529	-45.345	0.058 \pm 0.001	150.2 \pm 1.0	16.2 \pm 0.7	13.4 \pm 0.4	-3.5 \pm 0.7
* 287.545	-43.624	0.087 \pm 0.002	3.8 \pm 1.0	7.9 \pm 0.6	10.3 \pm 0.5	-14.0 \pm 0.6
287.675	-35.482	0.118 \pm 0.007	159.4 \pm 2.6	31.1 \pm 1.6	12.4 \pm 1.1	-1.3 \pm 1.6
* 288.067	-45.870	0.111 \pm 0.002	146.6 \pm 1.2	9.0 \pm 0.7	0.9 \pm 0.8	-9.8 \pm 0.7
288.421	-37.393	0.067 \pm 0.003	164.6 \pm 2.3	-30.1 \pm 1.3	4.7 \pm 2.2	-59.6 \pm 1.3
288.484	-33.311	0.139 \pm 0.008	111.5 \pm 4.0	18.8 \pm 2.1	2.3 \pm 1.9	-15.9 \pm 2.2
288.589	-39.501	0.017 \pm 0.001	81.8 \pm 3.1	-0.2 \pm 1.8	2.9 \pm 2.2	-26.9 \pm 1.8
* 288.627	-41.413	0.084 \pm 0.004	133.5 \pm 2.1	17.5 \pm 2.0	28.2 \pm 1.0	-6.8 \pm 2.0
* 288.715	-40.820	0.964 \pm 0.027	10.7 \pm 4.8	22.5 \pm 4.3	32.5 \pm 1.1	-2.5 \pm 4.3
289.090	-39.342	0.127 \pm 0.006	49.2 \pm 2.1	13.8 \pm 1.2	8.2 \pm 1.2	-12.9 \pm 1.2
289.105	-33.833	0.048 \pm 0.004	30.0 \pm 3.7	24.6 \pm 2.6	15.7 \pm 1.9	-9.1 \pm 2.6
* 289.145	-44.971	0.080 \pm 0.007	174.8 \pm 4.8	48.2 \pm 4.1	18.4 \pm 2.1	28.8 \pm 4.1
289.161	-32.618	0.112 \pm 0.007	28.2 \pm 2.5	37.1 \pm 1.3	6.9 \pm 1.8	1.8 \pm 1.3
289.167	-32.625	0.111 \pm 0.006	35.3 \pm 2.4	34.9 \pm 1.4	10.2 \pm 1.3	-0.3 \pm 1.4
* 289.253	-45.636	0.168 \pm 0.018	83.4 \pm 4.4	9.0 \pm 2.7	11.1 \pm 2.1	-9.5 \pm 2.7
289.395	-32.792	0.132 \pm 0.006	120.5 \pm 1.8	32.1 \pm 1.0	5.8 \pm 1.6	-2.8 \pm 1.0
289.478	-39.488	0.037 \pm 0.002	170.7 \pm 2.9	4.4 \pm 1.5	2.9 \pm 2.2	-21.9 \pm 1.5

Table B1 – continued

l ($^{\circ}$)	b ($^{\circ}$)	p_0 (per cent)	Ψ_0 ($^{\circ}$)	ϕ_{raw} (rad m^{-2})	σ_{ϕ} (rad m^{-2})	ϕ_{corr} (rad m^{-2})
* 290.012	-41.763	0.051 \pm 0.001	169.3 \pm 0.7	7.9 \pm 0.4	10.5 \pm 0.3	-15.2 \pm 0.4
290.038	-39.129	0.044 \pm 0.004	57.1 \pm 3.4	4.9 \pm 2.6	19.9 \pm 1.5	-21.5 \pm 2.6
290.434	-36.118	0.165 \pm 0.001	43.0 \pm 0.5	23.5 \pm 0.3	0.3 \pm 0.3	-6.6 \pm 0.3
290.710	-38.878	0.170 \pm 0.017	111.1 \pm 4.0	14.8 \pm 2.3	12.8 \pm 1.7	-11.6 \pm 2.3
290.754	-38.330	0.038 \pm 0.001	21.3 \pm 1.5	21.6 \pm 0.8	2.4 \pm 1.7	-5.5 \pm 0.8
290.754	-35.106	0.074 \pm 0.001	107.8 \pm 1.1	4.2 \pm 0.6	0.6 \pm 0.6	-27.0 \pm 0.6
290.852	-32.259	0.051 \pm 0.002	172.8 \pm 1.6	40.2 \pm 0.8	6.7 \pm 1.1	5.4 \pm 0.8
* 290.958	-45.418	0.028 \pm 0.003	49.3 \pm 4.6	13.1 \pm 4.4	22.7 \pm 1.9	-4.8 \pm 4.4
290.961	-40.949	0.130 \pm 0.006	158.0 \pm 2.1	3.7 \pm 1.1	5.7 \pm 1.8	-19.9 \pm 1.1
290.986	-36.119	0.058 \pm 0.002	35.5 \pm 1.3	2.0 \pm 0.8	11.0 \pm 0.6	-27.8 \pm 0.8
291.157	-38.395	0.127 \pm 0.001	3.6 \pm 0.4	11.4 \pm 0.2	1.6 \pm 0.9	-15.4 \pm 0.3
291.468	-44.002	0.536 \pm 0.026	90.4 \pm 3.2	6.2 \pm 1.9	2.3 \pm 1.9	-13.2 \pm 1.9
291.492	-41.596	0.033 \pm 0.001	179.4 \pm 1.2	11.3 \pm 0.6	5.6 \pm 0.9	-11.3 \pm 0.6
291.508	-40.460	0.199 \pm 0.005	54.3 \pm 1.3	0.3 \pm 0.8	11.8 \pm 0.5	-23.7 \pm 0.8
291.652	-34.097	0.079 \pm 0.006	175.1 \pm 5.0	11.5 \pm 5.7	39.8 \pm 2.0	-20.6 \pm 5.7
291.674	-34.518	0.044 \pm 0.001	48.1 \pm 0.6	23.2 \pm 0.4	0.5 \pm 0.5	-8.4 \pm 0.4
291.778	-40.785	0.059 \pm 0.007	37.7 \pm 4.6	6.0 \pm 3.6	21.1 \pm 2.2	-17.4 \pm 3.6
291.865	-37.269	0.134 \pm 0.010	91.7 \pm 2.9	37.7 \pm 2.6	26.7 \pm 1.5	9.9 \pm 2.6
291.865	-37.269	0.114 \pm 0.008	95.2 \pm 2.8	33.8 \pm 2.3	27.6 \pm 1.3	5.9 \pm 2.3
291.881	-35.986	0.075 \pm 0.004	58.2 \pm 2.3	12.9 \pm 1.2	6.7 \pm 1.6	-16.6 \pm 1.2
* 291.954	-31.336	0.161 \pm 0.006	87.6 \pm 2.0	38.1 \pm 1.4	19.3 \pm 0.7	2.7 \pm 1.4
292.041	-43.665	0.193 \pm 0.010	50.6 \pm 4.5	9.3 \pm 2.4	1.4 \pm 1.2	-10.3 \pm 2.4
292.082	-42.348	0.186 \pm 0.012	26.8 \pm 2.7	12.2 \pm 1.4	7.7 \pm 1.8	-9.0 \pm 1.4
292.397	-42.859	0.128 \pm 0.007	103.2 \pm 3.5	-1.2 \pm 1.8	3.0 \pm 2.2	-21.7 \pm 1.8
292.473	-37.351	0.018 \pm 0.000	129.6 \pm 1.1	23.6 \pm 0.6	3.9 \pm 1.5	-3.9 \pm 0.6
292.550	-41.942	0.545 \pm 0.026	65.5 \pm 2.1	9.1 \pm 1.3	13.2 \pm 0.8	-12.5 \pm 1.3
292.652	-44.236	0.293 \pm 0.003	19.9 \pm 0.4	18.2 \pm 0.3	15.8 \pm 0.2	-0.4 \pm 0.4
292.741	-43.379	0.235 \pm 0.019	87.9 \pm 3.5	15.9 \pm 1.8	7.3 \pm 2.3	-3.7 \pm 1.8
* 292.833	-31.083	0.081 \pm 0.003	35.4 \pm 2.4	40.7 \pm 1.2	2.1 \pm 1.7	5.4 \pm 1.2
292.868	-41.315	0.320 \pm 0.034	175.4 \pm 4.6	4.4 \pm 2.8	7.9 \pm 3.5	-17.8 \pm 2.8
293.055	-39.953	0.146 \pm 0.012	64.2 \pm 3.7	-0.8 \pm 2.0	10.2 \pm 1.6	-24.7 \pm 2.0
293.088	-43.924	0.363 \pm 0.033	78.7 \pm 5.0	27.7 \pm 3.9	8.3 \pm 4.9	9.0 \pm 3.9
293.148	-41.268	0.089 \pm 0.002	13.7 \pm 1.5	-20.0 \pm 1.1	19.4 \pm 0.5	-42.1 \pm 1.1
293.471	-42.911	0.178 \pm 0.007	122.4 \pm 1.7	10.0 \pm 1.0	11.2 \pm 0.7	-9.8 \pm 1.0
293.509	-43.572	0.317 \pm 0.014	80.4 \pm 1.9	20.6 \pm 1.1	7.9 \pm 1.2	1.6 \pm 1.1
293.635	-44.131	0.229 \pm 0.013	25.0 \pm 2.4	31.3 \pm 1.5	16.1 \pm 0.9	13.1 \pm 1.5
293.737	-42.296	0.284 \pm 0.018	130.6 \pm 2.8	10.9 \pm 1.6	6.9 \pm 2.2	-9.6 \pm 1.6
293.807	-41.469	0.129 \pm 0.012	61.1 \pm 3.9	13.5 \pm 2.3	11.8 \pm 1.7	-8.0 \pm 2.3
293.819	-41.466	0.154 \pm 0.012	71.9 \pm 3.3	6.4 \pm 2.0	14.2 \pm 1.4	-15.1 \pm 2.0
* 293.851	-31.371	0.035 \pm 0.000	98.3 \pm 0.3	43.8 \pm 0.2	0.1 \pm 0.1	9.4 \pm 0.2
293.907	-39.279	0.385 \pm 0.021	53.0 \pm 4.5	17.9 \pm 2.1	2.4 \pm 2.0	-6.3 \pm 2.1
294.006	-44.374	0.200 \pm 0.005	4.1 \pm 1.2	8.6 \pm 0.6	3.9 \pm 1.5	-9.1 \pm 0.7
294.192	-43.969	0.234 \pm 0.019	156.9 \pm 3.5	17.8 \pm 1.9	6.4 \pm 2.8	-0.3 \pm 1.9
294.271	-44.918	0.095 \pm 0.001	158.1 \pm 0.5	-8.4 \pm 0.3	3.1 \pm 0.7	-25.2 \pm 0.4
294.385	-44.286	0.084 \pm 0.003	93.6 \pm 1.7	22.4 \pm 1.0	9.5 \pm 0.9	4.7 \pm 1.0
* 294.480	-31.056	0.132 \pm 0.002	165.5 \pm 1.0	42.5 \pm 0.5	0.9 \pm 0.8	8.0 \pm 0.5
294.525	-40.882	0.081 \pm 0.004	99.8 \pm 1.9	2.4 \pm 1.1	8.3 \pm 1.2	-19.5 \pm 1.1
294.530	-42.244	0.347 \pm 0.024	12.2 \pm 3.6	18.4 \pm 1.8	4.3 \pm 2.7	-1.7 \pm 1.8
294.535	-40.875	0.063 \pm 0.002	95.2 \pm 1.5	4.9 \pm 0.9	5.7 \pm 1.4	-17.0 \pm 1.0
294.536	-43.701	0.454 \pm 0.019	53.2 \pm 1.8	21.9 \pm 0.9	7.9 \pm 1.0	3.6 \pm 0.9
294.548	-40.966	0.107 \pm 0.007	21.2 \pm 3.1	4.7 \pm 1.6	5.0 \pm 2.4	-17.1 \pm 1.6
294.593	-40.497	0.231 \pm 0.007	136.5 \pm 1.8	12.4 \pm 1.0	2.9 \pm 1.9	-10.0 \pm 1.0
294.608	-42.908	0.060 \pm 0.007	161.1 \pm 4.2	12.2 \pm 2.5	8.7 \pm 3.4	-7.1 \pm 2.5
294.713	-43.612	0.614 \pm 0.039	79.3 \pm 2.7	18.6 \pm 1.6	14.7 \pm 0.9	0.2 \pm 1.6
294.884	-44.122	0.052 \pm 0.001	73.1 \pm 1.5	26.3 \pm 0.8	2.8 \pm 1.8	8.8 \pm 0.9
294.912	-40.650	0.027 \pm 0.002	21.5 \pm 5.5	18.3 \pm 2.9	4.0 \pm 2.9	-3.7 \pm 2.9
294.930	-42.244	0.065 \pm 0.002	157.6 \pm 2.6	21.7 \pm 1.4	1.8 \pm 1.5	1.8 \pm 1.4
294.995	-41.978	0.051 \pm 0.002	49.8 \pm 1.7	25.5 \pm 1.1	9.5 \pm 1.1	5.3 \pm 1.1
295.126	-41.492	0.057 \pm 0.003	109.3 \pm 2.0	35.5 \pm 1.1	5.1 \pm 1.9	14.7 \pm 1.1
295.202	-41.984	0.167 \pm 0.014	86.4 \pm 3.5	28.6 \pm 2.0	9.8 \pm 1.8	8.4 \pm 2.0
295.226	-42.823	0.294 \pm 0.015	92.6 \pm 2.2	-9.5 \pm 1.2	11.6 \pm 1.0	-28.6 \pm 1.3

Table B1 – *continued*

l ($^{\circ}$)	b ($^{\circ}$)	p_0 (per cent)	Ψ_0 ($^{\circ}$)	ϕ_{raw} (rad m^{-2})	σ_{ϕ} (rad m^{-2})	ϕ_{corr} (rad m^{-2})
295.361	−40.387	0.121 ± 0.007	54.4 ± 4.1	25.0 ± 2.2	2.6 ± 2.2	2.9 ± 2.2
295.367	−40.785	0.322 ± 0.041	147.4 ± 10.5	22.9 ± 5.1	2.9 ± 2.6	1.3 ± 5.1
295.507	−42.082	0.162 ± 0.009	86.8 ± 2.2	27.7 ± 1.2	7.2 ± 1.7	7.8 ± 1.2
295.524	−40.997	0.062 ± 0.006	85.3 ± 6.4	11.5 ± 3.1	3.7 ± 2.9	−9.8 ± 3.1
295.675	−44.323	0.077 ± 0.007	173.7 ± 5.5	10.3 ± 2.8	5.3 ± 3.0	−6.6 ± 2.8
* 295.688	−34.852	0.130 ± 0.003	9.9 ± 1.4	30.8 ± 0.9	5.7 ± 1.1	1.8 ± 0.9
295.733	−42.013	0.142 ± 0.010	105.0 ± 3.4	34.9 ± 2.0	5.8 ± 2.7	15.1 ± 2.0
295.814	−43.417	0.274 ± 0.030	174.5 ± 5.0	21.8 ± 2.8	8.3 ± 3.2	3.8 ± 2.8
295.881	−43.599	0.022 ± 0.001	72.5 ± 2.7	42.0 ± 1.8	15.8 ± 0.9	24.3 ± 1.8
295.893	−42.505	0.205 ± 0.011	48.4 ± 2.6	19.3 ± 1.6	15.5 ± 0.9	0.2 ± 1.6
295.925	−45.474	0.473 ± 0.037	85.4 ± 3.6	−5.1 ± 2.3	17.5 ± 1.1	−20.4 ± 2.3
295.925	−43.169	0.209 ± 0.014	73.3 ± 3.2	18.8 ± 1.9	9.9 ± 1.6	0.5 ± 1.9
295.956	−43.659	0.044 ± 0.001	37.3 ± 1.4	23.1 ± 0.8	6.6 ± 1.1	5.4 ± 0.9
295.963	−43.664	0.054 ± 0.002	39.0 ± 1.4	21.7 ± 1.0	12.1 ± 0.7	4.1 ± 1.0
295.986	−42.955	0.325 ± 0.054	106.3 ± 7.2	6.3 ± 3.8	9.5 ± 3.9	−12.2 ± 3.8
296.022	−42.163	0.070 ± 0.004	50.4 ± 3.2	2.5 ± 3.4	31.1 ± 1.4	−17.0 ± 3.4
296.491	−40.813	0.226 ± 0.011	157.0 ± 3.4	4.5 ± 1.9	1.6 ± 1.4	−16.4 ± 1.9
296.659	−45.653	0.101 ± 0.004	70.9 ± 1.6	−13.1 ± 1.0	10.6 ± 0.8	−27.8 ± 1.0
296.704	−43.455	0.057 ± 0.006	168.7 ± 3.9	6.0 ± 2.4	9.2 ± 2.7	−11.5 ± 2.4
296.719	−40.842	0.194 ± 0.004	55.0 ± 0.8	4.7 ± 0.5	6.4 ± 0.6	−16.1 ± 0.5
296.882	−40.719	0.237 ± 0.028	52.9 ± 4.7	−7.1 ± 2.3	8.5 ± 2.8	−28.0 ± 2.3
* 296.933	−33.860	0.137 ± 0.011	6.3 ± 4.0	27.3 ± 3.2	20.9 ± 1.4	−2.4 ± 3.2
296.997	−40.395	0.063 ± 0.002	68.3 ± 1.4	1.3 ± 0.7	2.6 ± 1.6	−19.9 ± 0.7
297.068	−41.714	0.043 ± 0.002	125.5 ± 2.3	16.8 ± 1.3	11.3 ± 0.9	−2.8 ± 1.3
297.070	−41.711	0.042 ± 0.002	114.8 ± 2.7	20.2 ± 1.5	10.8 ± 1.0	0.7 ± 1.5
297.070	−41.257	0.063 ± 0.003	145.0 ± 3.4	18.4 ± 1.8	1.7 ± 1.5	−1.7 ± 1.8
297.257	−41.186	0.300 ± 0.013	132.7 ± 2.8	17.6 ± 1.8	2.8 ± 2.2	−2.5 ± 1.8
297.277	−42.705	0.628 ± 0.056	9.5 ± 6.1	11.9 ± 3.7	4.6 ± 3.9	−6.3 ± 3.7
297.344	−43.704	0.233 ± 0.019	67.3 ± 3.4	−18.4 ± 2.1	8.6 ± 2.2	−35.2 ± 2.1
297.476	−41.179	0.766 ± 0.128	150.1 ± 31.9	−21.9 ± 12.1	5.1 ± 3.4	−41.9 ± 12.1
297.624	−44.093	0.170 ± 0.017	76.8 ± 3.9	−11.0 ± 2.1	8.3 ± 2.6	−27.2 ± 2.1
297.758	−44.158	0.097 ± 0.007	33.4 ± 3.2	−14.2 ± 2.1	15.5 ± 1.4	−30.3 ± 2.1
* 298.169	−35.710	0.127 ± 0.003	114.9 ± 1.1	50.7 ± 0.6	4.2 ± 1.3	24.0 ± 0.6
298.209	−42.078	0.177 ± 0.014	15.6 ± 4.4	−15.7 ± 2.4	3.7 ± 2.7	−34.2 ± 2.4
* 298.247	−33.110	0.125 ± 0.010	25.8 ± 3.1	38.3 ± 1.8	7.1 ± 2.6	8.4 ± 1.8
298.251	−42.116	0.433 ± 0.026	11.8 ± 2.7	−46.6 ± 1.6	7.7 ± 1.9	−65.0 ± 1.6
298.381	−43.653	0.050 ± 0.003	144.7 ± 3.3	7.3 ± 1.8	3.1 ± 2.3	−9.1 ± 1.8
* 298.994	−36.643	0.031 ± 0.002	135.7 ± 2.8	18.3 ± 2.0	15.9 ± 1.1	−6.8 ± 2.0
* 299.493	−30.568	0.021 ± 0.001	2.7 ± 2.0	24.7 ± 1.2	11.6 ± 0.8	−7.9 ± 1.2
* 299.727	−33.232	0.019 ± 0.001	178.4 ± 2.5	36.5 ± 1.5	13.4 ± 0.9	7.5 ± 1.5
299.815	−41.686	0.190 ± 0.003	23.6 ± 1.4	38.4 ± 1.0	0.9 ± 0.8	20.2 ± 1.0
* 300.101	−32.957	0.107 ± 0.002	42.9 ± 1.5	19.3 ± 0.8	1.3 ± 1.1	−9.9 ± 0.8
* 300.176	−37.791	0.083 ± 0.004	156.5 ± 3.2	25.4 ± 1.7	1.8 ± 1.5	2.4 ± 1.7
300.260	−41.713	0.022 ± 0.000	32.5 ± 1.2	21.2 ± 0.6	1.6 ± 1.2	3.3 ± 0.7
* 300.546	−34.430	0.210 ± 0.004	35.0 ± 1.6	19.8 ± 0.9	0.8 ± 0.7	−7.3 ± 0.9
* 300.682	−38.509	0.098 ± 0.004	124.7 ± 1.6	31.4 ± 1.0	11.6 ± 0.7	9.6 ± 1.0
* 300.712	−31.222	0.023 ± 0.001	26.9 ± 1.6	12.4 ± 0.9	5.8 ± 1.3	−18.7 ± 0.9
* 300.977	−35.998	0.096 ± 0.009	60.8 ± 4.0	33.9 ± 3.0	20.0 ± 1.6	9.1 ± 2.9
* 301.077	−37.858	0.378 ± 0.032	144.6 ± 3.6	28.0 ± 1.7	7.7 ± 1.8	5.5 ± 1.7
* 301.241	−35.873	0.042 ± 0.002	157.0 ± 2.3	8.7 ± 1.2	11.3 ± 0.9	−16.2 ± 1.2
* 301.463	−32.596	0.089 ± 0.003	37.9 ± 1.3	28.7 ± 0.8	10.9 ± 0.7	−0.2 ± 0.8
* 302.602	−38.463	0.059 ± 0.002	177.2 ± 1.8	10.8 ± 1.0	7.1 ± 1.2	−10.1 ± 1.0
* 304.115	−35.992	0.041 ± 0.003	138.4 ± 4.0	48.2 ± 2.8	15.4 ± 1.5	25.0 ± 2.8

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