

Proposal to Detect Dark Matter using Axionic Topological Antiferromagnets

David J. E. Marsh^{1,*}, Kin Chung Fong,² Erik W. Lentz,¹ Libor Šmejkal,^{3,4,5} and Mazhar N. Ali⁶

¹*Institut für Astrophysik, Georg-August Universität, Friedrich-Hund-Platz 1, D-37077 Göttingen, Germany*

²*Raytheon BBN Technologies, Quantum Engineering and Computing, Cambridge, Massachusetts 02138, USA*

³*Institut für Physik, Johannes Gutenberg Universität Mainz, D-55099 Mainz, Germany*

⁴*Institute of Physics, Academy of Sciences of the Czech Republic, Cukrovarnická 10, 162 53 Praha 6 Czech Republic*

⁵*Faculty of Mathematics and Physics, Charles University in Prague, Ke Karlovu 3, 121 16 Prague 2, Czech Republic*

⁶*Max Planck Institute of Microstructure Physics, Weinberg 2, 06120 Halle (Saale), Germany*



(Received 13 August 2018; published 17 September 2019; corrected 16 October 2019)

Antiferromagnetically doped topological insulators (ATI) are among the candidates to host dynamical axion fields and axion polaritons, weakly interacting quasiparticles that are analogous to the dark axion, a long sought after candidate dark matter particle. Here we demonstrate that using the axion quasiparticle antiferromagnetic resonance in ATIs in conjunction with low-noise methods of detecting THz photons presents a viable route to detect axion dark matter with a mass of 0.7 to 3.5 meV, a range currently inaccessible to other dark matter detection experiments and proposals. The benefits of this method at high frequency are the tunability of the resonance with applied magnetic field, and the use of ATI samples with volumes much larger than 1 mm³.

DOI: [10.1103/PhysRevLett.123.121601](https://doi.org/10.1103/PhysRevLett.123.121601)

Astrophysical and cosmological observations provide strong evidence for the existence of nonbaryonic dark matter (DM) [1–4]. Among possible candidates are dark axions (DA) [5–11], which are hypothetical particles [12–14] suggested to solve the charge-parity (\mathcal{CP}) problem in quantum chromodynamics (QCD) [15]. Searching for the DA is challenging due to its weak coupling to ordinary matter (e.g., photons). For DA masses $m_a \lesssim 0.2$ eV the local DA field, θ_D , can be described as a classical coherent state. The local DM density is then $\rho_{\text{DM}} = |\theta_D(t)|^2 m_a^2 f_a^2 / 2$, where m_a and f_a are the unknown axion mass and “decay constant,” and the measured value is $\rho_{\text{DM}} \approx 0.4 \text{ GeV cm}^{-3}$ [3]. The DA field oscillates in time, with a frequency dominated by the rest energy, $m_a c^2$, and an intrinsic width set by the galactic velocity dispersion, $\sigma_v \approx 230 \text{ km s}^{-1} \Rightarrow \Delta\omega_a / \omega_a = \sigma_v^2 / c^2 \approx 10^{-6}$. The QCD axion mass can be computed in chiral perturbation theory or on the lattice, and is given by $m_a = 0.6 \text{ meV} (10^{10} \text{ GeV} / f_a)$ [13,14,16] (we use units $\hbar = c = 1$ if not stated otherwise). The central frequency is $\nu = 0.25 (m_a / \text{meV}) \text{ THz}$.

Only one DM search, the axion dark matter experiment (ADMX) [17,18], has made a significant constraint on the QCD axion parameter space predicted by the Kim-Shifman-Vainshtein-Zhakarov (KSVZ) [19,20] and

Dine-Fischler-Srednicki-Zhitnitsky (DFSZ) [21,22] models. The QCD axion mass can span $10^{-12} \lesssim m_a \lesssim 10^{-2}$ eV, satisfying astrophysical constraints on the couplings [23,24] with f_a less than the Planck scale. There are hints, however, pointing to the meV range, particularly for DFSZ-type models [25,26]. Furthermore, constraints from the CERN axion solar telescope (CAST) [27] and the any light particle search [28], combined with the prediction of KSVZ and DFSZ models provide a target range for the axion-photon coupling $10^{-13} \text{ GeV}^{-1} < g_\gamma < 10^{-10} \text{ GeV}^{-1}$ for $m_a = 1$ meV. In the meV range, axion DM searches also overlap with searches for spin-dependent forces [29,30].

The power output from the axion-induced electric field [Fig. 1(a)] is:

$$P_0 = \frac{1}{2} E_{\text{eff}}^2 V_{\text{eff}} \omega_a = g_\gamma^2 B_0^2 \frac{\rho_{\text{DM}}}{m_a^2} V_{\text{eff}} \omega_a. \quad (1)$$

Taking the effective volume $V_{\text{eff}} \approx (2\pi/m_a)^3$ from the vacuum dispersion relation, $m_a = 1$ meV, $g_\gamma = 10^{-10} \text{ GeV}^{-1}$, and $B_0 = 1$ T, gives $P_0 = 10^{-27}$ W. Detecting the axion requires amplifying this power. Methods to amplify the signal include resonance in a microwave cavity [17,18,31,32], ferromagnetic resonance [33], coherent enhancement [34], and many others [33,35–41]. In particular, a resonant cavity haloscope method at mm wavelengths enhances the signal by the quality factor, Q , but suffers from small effective volume, since the resonance requires $V \sim (2\pi/\omega_a)^3$. The highest frequency operating cavity haloscope is ORGAN, at 0.1 meV [32], while the MADMAX dielectric haloscope projects maximum

Published by the American Physical Society under the terms of the [Creative Commons Attribution 4.0 International license](https://creativecommons.org/licenses/by/4.0/). Further distribution of this work must maintain attribution to the author(s) and the published article's title, journal citation, and DOI. Funded by SCOAP³.

frequencies of 0.5 meV [34]. In this Letter we propose an alternative method that combines THz resonant enhancement and volume increase facilitated by axion-photon conversion inside a topological axion insulator antiferromagnet. To estimate the signal strength we use antiferromagnetically Fe-doped Bi_2Se_3 as a realistic possibility, demonstrating that it fits our general requirements through symmetry analogy to the Fu-Kane-Mele-Hubbard model of antiferromagnetic diamond.

Axionic degrees of freedom are predicted to materialize as quasiparticles, θ_Q , in magnetically doped topological insulators (TIs) [45], Cr_2O_3 ($\theta_Q = \pi/36$) [46,47], $\alpha\text{-Fe}_2\text{O}_3$ [48] with a corundum structure, spinels [49], and magnetic TI heterostructures [50]. The signatures of the topological magnetoelectric effect, also known as *static* axion electrodynamics, were recently reported as quantized magneto-optical effects in TIs [51–53], and quantized magneto and electrical resistance changes in artificial antiferromagnetic heterostructures of magnetically doped TIs [54–56]. Finally, dynamical axion quasiparticles (AQ) in the form of magnetic fluctuations were predicted in magnetically doped TIs (MTI) [45], spin-orbit coupled Mott insulators [57], and in MTI superlattices [50].

We propose to use AOs in antiferromagnetically doped TIs (ATI) to detect DAs. The conversion process of DAs to visible photons is shown in Fig. 1(b). Antiferromagnets provide the correct THz frequency range owing to the resonance frequency exchange enhancement $\omega \sim \sqrt{(2H_E + H_A)H_A}$ (H_E, H_A are exchange and anisotropy fields respectively). Inside the ATI, AOs mix with the electric field E and generate (quasiparticle) axion polaritons [AP, see Figs. 1(a) and 1(b)], ϕ_{\pm} [45] (see also Ref. [58]). When $\omega_{\pm}(k, B_0) = \omega_a$, the conversion process is resonantly enhanced by $Q = \omega/\Gamma$, where Γ is the polariton damping (width). If the allowed values of k are restricted by the geometry, then the lowest value of $k \sim 1/L$ can facilitate resonant conversion of DAs to APs in volumes much larger than $(c/\text{THz})^3$. The combination of Q and V allows the signal power to be greatly enhanced compared to P_0 . The APs convert into propagating photons due to the boundary conditions (BCs) [44], and can be detected. As we will now show, this detection strategy gives access to a unique part of DA parameter space.

We begin by defining an axionic field θ_i by the coupling to the electromagnetic Chern-Simons (CS) term generated by the loop [Fig. 1(a)]:

$$S_{\text{CS}} = \sum_{i=\text{D,Q}} \frac{\alpha}{\pi} C_i \int d^4x \theta_i \mathbf{E} \cdot \mathbf{B}, \quad (2)$$

where \mathbf{E}, \mathbf{B} are electric and magnetic fields. θ_D , is a pseudoscalar pseudo-Goldstone boson with a nonvanishing electromagnetic chiral anomaly [12–14,42,43]. The coupling C_i is dimensionless: the dimensionful axion-photon coupling is defined by $g_\gamma = C_i \alpha / 2\pi f_i$. C_D is a

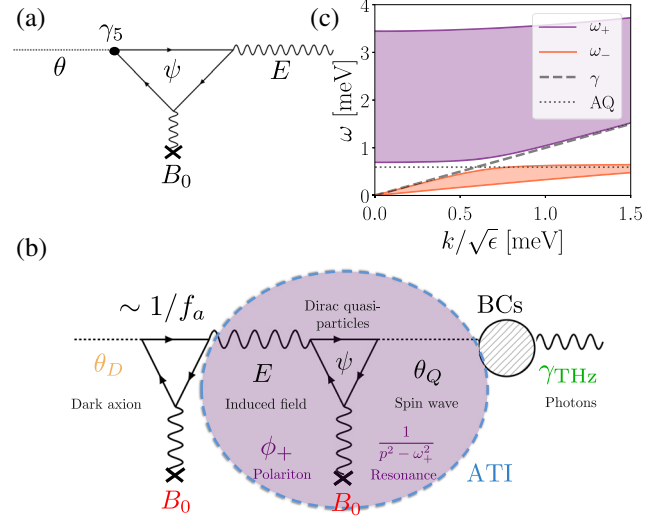


FIG. 1. (a) The chiral anomaly [42,43]. θ is a pseudoscalar chirally coupled to charged Dirac fermions, ψ . With applied B_0 , θ mixes with E leading to the existence of axion polaritons, ϕ_{\pm} , in the case of θ_Q and the production of photons in the case of θ_D . (b) Resonant enhancement of DA-photon conversion. Colored text refers to Fig. 3. Inside the ATI, the DA couples to the mixed states ϕ_{\pm} shown in the shaded circle. Conversion is resonantly enhanced when $p^2 = \omega_a^2 = \omega_{\pm}(k, B_0)^2$, represented by the polariton propagator. At the ATI dielectric boundary, polaritons convert to propagating photons, due to boundary conditions (BCs) [44] represented here by the vertex. (c) The axion-polariton dispersion relation for $\omega_{\pm}(k, B_0)$ [45]. Scanning the applied B_0 field tunes $\omega_{\pm}(k=0)$ in the range 0.7 to 3.5 meV and scans the resonance.

model-dependent constant taking the values $C_{\text{KSVZ}} = -1.92$ and $C_{\text{DFSZ}} = 0.75$, and $f_i = f_a$. For the AQ, we define $C_Q = 1$.

Other DA couplings to ordinary matter [10,59,60] could also affect the ATI. Nuclear spin couplings lead to resonance at the Larmor frequency, which with $B \lesssim 20$ T gives $\nu \lesssim 100$ MHz [39,60], far below the DA frequency at 1 meV. The axion-electron coupling induces DA absorption in Dirac semimetals [61]. The parameter space with significant absorption, however, is excluded by astrophysical constraints. Thus we neglect the direct nuclear and electron DA couplings.

The criteria for generating AOs in condensed matter as suggested by Wilczek are [62] (i) effective action in the form of Eq. (2), (ii) realization of the Dirac equation for electrons, and (iii) tuneable Dirac masses.

Criterion (i) can be met in general in magnetoelectric materials with nonzero diagonal components of the magnetoelectric polarizability tensor $\alpha_{ij} = (\partial M_j / \partial E_i)_{\mathbf{B}=0} = (\partial P_i / \partial B_j)_{\mathbf{E}=0}$, where M, P are magnetization and electric polarization. Since θ_Q is odd under spatial inversion \mathcal{P} and time reversal \mathcal{T} , and the physical observables $\sim e^{iS/\hbar}$ (where S is the action) are defined modulo 2π , the CS term can be nonzero in (a) magnetoelectric materials with a magnetic

point groups with broken \mathcal{P} , and broken \mathcal{T} where θ is nonquantized, (b) $\theta = \pi$ can be taken as a defining property of \mathcal{T} -invariant TIs [46,51].

Criterion (ii) can be realized in Dirac quasiparticle materials such as TIs where the simultaneous presence of \mathcal{P} and \mathcal{T} symmetries protects the Kramers double degeneracy of the bulk Dirac bands, while at the surfaces realize \mathcal{T} protected 2D Dirac quasiparticle helical states [63]. To satisfy (iii) and generate dynamical axion fields, gradients of θ need also be generated dynamically, one possibility being magnetic fluctuations [45,50,57]. In such a case, θ_Q is the pseudoscalar component of the spin wave.

To simultaneously satisfy all three criteria for AQs we identify Dirac quasiparticle antiferromagnets as suitable candidates [64–66]. We consider a Dirac antiferromagnetic insulator with \mathcal{P} and \mathcal{T} symmetry broken and thus magnetoelectric point group, but importantly the combination \mathcal{PT} preserved, with a generic electronic Dirac Hamiltonian $H(\mathbf{k}) = \sum_{i=1,\dots,5} A_i(\mathbf{k})\gamma_i$, where γ_i are Dirac matrices and $A_i(\mathbf{k})$ parametrize the band structure. The antiferromagnetic coupling couples in proper basis choice to γ_5 in the Dirac Hamiltonian [45,66].

As a particular realization we consider the antiferromagnetic Fu-Kane-Mele-Hubbard [67] model on the bipartite (orbital degree of freedom τ) diamond lattice with two spins per lattice site σ [66,68]. The antiferromagnetism breaks \mathcal{T} and \mathcal{P} , but preserves \mathcal{PT} as marked by the red ball in Fig. 2(a) and thus preserves the form of the Dirac Hamiltonian. The Hamiltonian with a Hubbard term treated on a mean-field level reads: $H = \lambda[\mathbf{A}(\mathbf{k}) - (U/\lambda)\mathbf{m}] \cdot \sigma\tau_z + t\text{Re}f(\mathbf{k})\tau_x + t\text{Im}f(\mathbf{k})\tau_y$, where the nearest neighbor hopping on the diamond lattice [cf., Fig. 2(a)] $f(\mathbf{k}) = \sum_{j=1,\dots,4} (t + \delta t_j) e^{ik \cdot \mathbf{d}_j}$ (\mathbf{d}_j being the four nearest neighbor

vectors), $A_x(\mathbf{k}, Um_x) = 4\sin(k_x/2)[\cos(k_y/2) - \cos(k_z/2)]$ plus cyclic permutations, U is Hubbard correlation strength, λ is strength of the spin-orbit coupling, and δt_j represent the renormalization of the hopping due the deformation of the AB bond. The AQ has a mean value given by $\theta_Q = \pi/2[1 + \text{sgn}(\delta t_1)] - \arctan(Um/\delta t_1)$ [66]. It was shown that the fluctuations in the Néel order parameter $\mathbf{L} \sim \mathbf{m}_A - \mathbf{m}_B$ (with axis along z) can be within the approximation $(U|\mathbf{m}|/\lambda) \ll 1$ related to dynamical fluctuations of θ_Q [45,66]:

$$\delta\theta_Q \sim \frac{2}{3} \sum_{i=x,y,z} Um_i. \quad (3)$$

The band structure of our model is shown in Fig. 2(a) for a realistic range of effective exchange coupling $Um \sim 0-0.40$ and illustrates the tuning of the Dirac bands with a Dirac point shifted slightly off the X ($\sim Um/2\lambda$) point due to the effect of antiferromagnetism. The AQ spin wave (SW) [69–71] dispersion on the diamond lattice is

$$\hbar\omega_{QA} \approx g\mu_B H_0 \pm \sqrt{[8SJf(0) + g\mu_B H_A]^2 - [8SJf(\mathbf{q})]^2}, \quad (4)$$

where $g \approx 1$ is the Landé factor, μ_B is the Bohr magneton, $H_E = 8SJ$, and \mathbf{q} is the spin wave vector. The AQ-SW tunes, in a first-order approximation, only the z -component of the $\mathbf{L} \propto \mathbf{M}_A - \mathbf{M}_B$ order parameter [45,66] (where $\mathbf{M}_{A,B}$ is the magnetization on A, B sublattices), which therefore tunes the Dirac mass as schematically illustrated by the shaded region in Fig. 2(a).

No antiferromagnetic bulk dynamical axionic insulator has yet been identified in the lab. Remarkably, however, our model has exactly the same magnetic point group, $\bar{3}'1m'$, as the mean-field medium of Fe-doped Bi_2Se_3 . This can be seen by deforming the face centered cubic primitive unit cell [Fig. 2(a)] along the [111] direction to produce the rhombohedral unit cell of tetradymite Bi_2Se_3 [Fig. 2(b)]. It can be shown that the antiferromagnetism couples to the same γ_5 matrix as in our model [45], and applying the Neumann principle gives the axion field favorable nonzero diagonal symmetric elements to α_{ij} , and leads to the analogous expression for the AQ-SW field [Eq. (3)].

The quadratic action for small fluctuations $|\theta_Q| < 1$ is given by:

$$S_{\text{AQ}} = \frac{f_Q^2}{2} \int d^4x [\dot{\theta}_Q^2 - (v_{Q,i} \partial_i \theta_Q)^2 - m_Q^2 \theta_Q^2], \quad (5)$$

where f_Q and \mathbf{v}_Q are the SW stiffness and velocity. Scanning $\omega_{\pm}(B_0)$ [see Fig. 1(c)] requires specifying $m_s(B_0)$ and $f_Q(B_0)$. For $(\text{Bi}_{1-x}\text{Fe}_x)_2\text{Se}_3$ using Eq. (4) with doping factor at 3.5% [72], exchange of 1 meV [73]

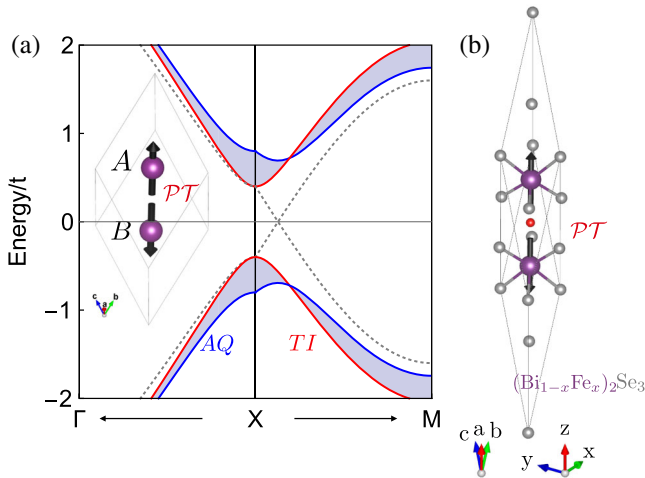


FIG. 2. Axion spin wave Dirac quasiparticle antiferromagnets. (a) Band structure of tuneable Dirac quasiparticles. TI: $\lambda = 0.5$, $\delta t_1 = 0.4$, $Um = 0$. AQ: $\lambda = 0.5$, $\delta t_1 = 0.4$, $0 < Um < 0.25$. Inset: Antiferromagnetic diamond lattice with marked \mathcal{PT} symmetry. (b) Crystal of AF-doped $(\text{Bi}_{1-x}\text{Fe}_x)_2\text{Se}_3$ exhibits the same magnetic point group symmetry as (a).

and anisotropy of 16 meV [74], the spin wave mass is $m_Q = [0.12(B_0/2 \text{ T}) + 0.6] \text{ meV}$. From Ref. [45] we find $f_Q = 190 \text{ eV}$ at $B_0 = 2 \text{ T}$, and take $f_Q^2 \propto 1/m_Q$ from the $\delta\mathbf{L}$ kinetic term [50].

Including the usual Maxwell term, linearizing for small fluctuations in E and θ_Q in the presence of an applied magnetic field, B_0 , and external DA source, we find the system of equations derived from the action take the following form (see also Refs. [44,62,75]):

$$\begin{aligned} \epsilon \ddot{\mathbf{E}} - \nabla^2 \mathbf{E} + \frac{\alpha}{\pi} [\mathbf{B}_0 \ddot{\theta}_Q - \nabla(\nabla\theta_Q \cdot \mathbf{B}_0)] &= \mathbf{A} \cos \omega_a t, \\ \ddot{\theta}_Q - v_Q^2 \nabla^2 \theta_Q + m_Q^2 \theta_Q - \frac{\alpha}{4\pi^2 f_Q^2} \mathbf{B}_0 \cdot \mathbf{E} &= 0, \end{aligned} \quad (6)$$

where $\epsilon = \epsilon_r \epsilon_0$ is the TI dielectric constant.

The driving term $\mathbf{A} = 2\mathbf{B}_0 g_\gamma \sqrt{2\rho_{\text{DM}}}/m_a$ at leading order, and derives from Eq. (2) taking the DA as an external source, with θ_D fixed by ρ_{DM} . Neglecting the AQ dispersion compared to \mathbf{E} , we diagonalize Eq. (6) to find ϕ_\pm and $2\omega_\pm^2(k) = (k^2/\epsilon + m_Q^2 + b^2) \pm \sqrt{(k^2/\epsilon + m_Q^2 + b^2)^2 - 4k^2 m_Q^2/\epsilon}$, where $b^2 = \alpha^2 B_0^2 / 4\pi^3 \epsilon f_Q^2$ [45] [see Fig. 1(c)], and k is the Fourier conjugate of x . Dynamical AOs are required for the mixing: in the absence of derivatives, θ_Q and E decouple in Eq. (6). The presence of axion-polaritons can be verified using an inverse ‘‘light shining through a wall’’ [76] experiment (as described elsewhere [45]), which can measure $\omega_\pm(k, B_0)$ from the band gap.

DA-driven polariton waves in the ATI are a combination of \mathbf{L} , and \mathbf{E} . In the presence of \mathcal{T} breaking, the ATI surface states are gapped [45,77]. The DA-induced surface polariton E field thus leads to emission of photons from the surface of the ATI, just like a dielectric haloscope, or dish antenna [34,78]. If there is only one mode at a given ω then dielectric BCs are sufficient to compute the photon emission from polaritons at the boundary. We propose to detect the emitted photons by using a silicon lens to focus them onto a wide bandwidth single photon detector (SPD). A mirror placed behind the ATI coherently enhances the forward emission [44]. The concept is illustrated in Fig. 3.

The material boundary conditions and mirror restrict the allowed modes k in the ATI. The lowest lying mode, $k_0 \sim 1/L$ (where L is the system size), has the largest integrated volume, and we define $\omega_+(k_0, B_0) := \omega_+^0(B_0)$. The E -field power generated by resonant DA-photon conversion on resonance, $\omega_+^0(B_0) = \omega_a$, can be expressed as [79], $P_{\text{signal}} = (\omega/Q) \times (\text{energy stored})$:

$$P_{\text{signal}} = \frac{1}{2} \kappa f_+ Q_{\text{sys}} V_{\text{eff}} |E_0|^2 \omega_a, \quad (7)$$

where Q_{sys} is the loaded quality factor, κ is a coupling or form factor, and $f_+ = b^2/(\omega^2 + b^2)$ is a mode mixing

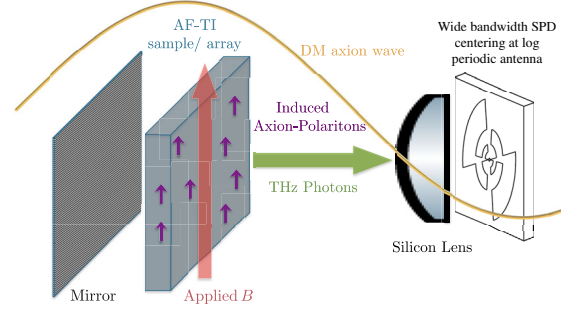


FIG. 3. Schematics of experimental concept. EM waves are emitted from all material surfaces perpendicular to B_0 . A mirror and silicon lens focus THz photons from ATI samples onto a single-photon detector located at the center of a log periodic antenna.

factor. We expect $\kappa \sim 1/\epsilon$; however this could be enhanced by $1/k^2$ at small $k \sim \sqrt{\epsilon} m_a$ due to resonant mixing.

The reference power, P_0 [Eq. (1)] is enhanced in Eq. (7) by two factors: first the quality factor of the AF resonance; second the effective volume can be far larger than $(2\pi/\omega)^3$. The volume amplification arises from the modified dispersion relation: the resonance is tuned by B_0 , and is independent of the ATI volume (the resonance scanning requires no precision THz mechanical motion at cryogenic temperatures). The mode mixing factor leads to a small suppression of power, and determines the optimal material via b . The coupling factor, κ , should be optimized in engineering of coatings, geometry, and material ϵ .

The effective Q_{sys} is due to the electric field enhancement inside the ATI due to the modified dispersion relation in Eq. (6). We assume $Q_{\text{sys}} = 10^5$, using THz AFMR measurements that report $Q \sim 10\text{--}100$ at $T \sim 4 \text{ K}$, reducing at lower T [80–84], and scaling $Q \propto T^{-3}$ [81,85–87] down to $\sim 100 \text{ mK}$ dilution refrigerator temperatures common in axion searches [18]. The value of Q and operating temperature will be key drivers in the final choice of material and experimental design.

The polariton in the ATI should be optimally coupled to the free space electromagnetic field at the surface for efficient photon measurement, and material losses due to Gilbert damping and phonon production [additional decay channels in Fig. 1(b)] should be of order the photon emission. We absorb into V_{eff} [see Eq. (1)] the relevant form factors, the effect of the ATI dielectric constant, and any boost factor, β^2 , arising from the geometry [44]. For $V_{\text{eff}} = 1 \text{ cm}^3$, $g_\gamma = 10^{-10} \text{ GeV}^{-1}$, $B_0 = 2 \text{ T}$ ($\omega_a = 0.8 \text{ meV}$, $\nu = 210 \text{ GHz}$), and $\kappa = 0.01$, the power is $5 \times 10^{-22} \text{ W}$: about one photon every 0.3 seconds.

We use SPD to estimate the measurement sensitivity because, at low temperatures and high frequencies, it is more advantageous than power detection [88]. While phase-insensitive linear amplifications are fundamentally limited by the standard quantum limit, SPD suffers no strict sensitivity limit. A high confidence detection requires the

dark count rate, Γ_d , of the detector to be smaller than the flux. We use $\Gamma_d = 0.001$ Hz, which has been demonstrated for the quantum dot detector in THz regime at 0.05 K [89]. A wider bandwidth, lower dark count SPD using graphene-based Josephson junction [90] has the potential to improve significantly the search for heavy dark axions in the future, including our proposal.

We propose to shield backgrounds by placing the entire apparatus in a cryostat, and then measure the baseline photon count at $B_0 = 0$. Measuring the dependence of the signal on B_0 and other features of the theoretical DA line shape (measured using a bandpass) [17,18,91] will allow candidate lines to be distinguished from signal.

The range of axion masses accessible to our technique depends on the scaling of material properties with B_0 . We take $1\text{ T} < B_0 < 10\text{ T}$ with stability $\delta B_0 = 10^{-3}\text{ T}$ over the volume, which has been demonstrated [18,92]. For the parameters of $(\text{Bi}_{1-x}\text{Fe}_x)_2\text{Se}_3$ given above and setting $\omega_+(k=0) = m_a$ we find $0.7\text{ meV} \leq m_a \leq 3.5\text{ meV}$ (the lower limit is approximately the $B_0 = 0$ spin wave mass). Other materials with different anisotropy field strengths can cover a wider range of masses.

Sensitivity to g_γ is computed setting the signal to noise ratio $\text{SNR} = 3$. We take the measurement time on a single frequency $\tau = 10^2\text{ s}$. The full range can be scanned in 6 months. The volume of any single, high quality sample of ATI is limited to be less than 1 cm^3 to achieve homogeneous doping [93]. The sensitivity is shown in Fig. 4 (stage-I).

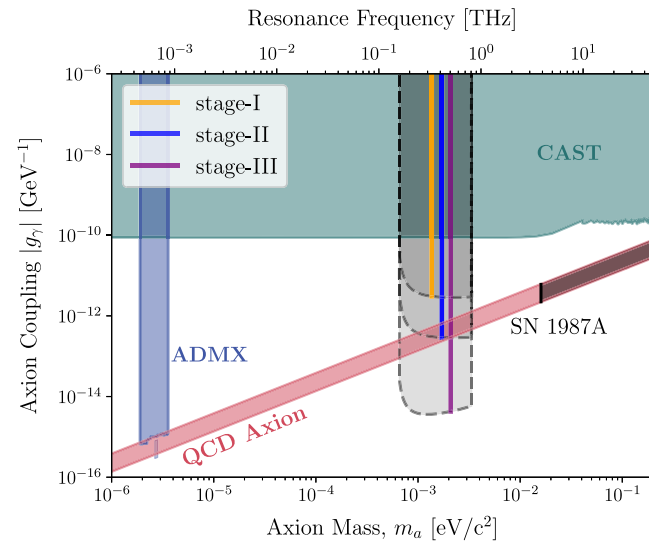


FIG. 4. Axion parameter space. Vertical lines show the projected sensitivity of our proposal using Fe-doped Bi_2Se_3 at $\sim 5\text{ T}$ applied field for 10^2 s integration time with dark count rate $\Gamma_d = 0.001$ Hz. Staged designs are described in the text. Gray shaded regions assume scanning $1\text{ T} \leq B_0 \leq 10\text{ T}$. The KSVZ and DFSZ axion models are shown as the red band. Existing exclusions from ADMX [17,18], CAST [27], and supernova 1987A [23] are shown as colored regions.

Using N ATI samples, either with a simple tiling and use of lenses, or with coherent addition [34], the gain in V_{eff} can increase linearly with N , with wide band response [44]. With $N = 100$ (a feasible total number for solid state synthesis [94]), the increased sensitivity is shown in Fig. 4 (stage-II).

A further increase in V_{eff} can be achieved by surrounding the ATI samples with a cavity with a volume, V_c . Long wavelength modes of the cavity E field can couple to high frequency AP modes resulting in a TM_{010} type [95] component to the AP, allowing $V_{\text{eff}} \approx V_c$ even with a small sample volume. In Fig. 4 (stage-III) we show the sensitivity benefit of a $V_{\text{eff}} = (0.1\lambda_{\text{dB}})^3 \approx 2000(1\text{ meV}/m_a)^3\text{ cm}^3$. The same stage-III sensitivity could be achieved if technology and investment allowed for fabrication of a very large volume of ATI.

In summary, we have shown that ATIs can host dynamical axionic quasiparticles which are resonantly driven in the presence of DAs with mass of order 1 meV and emit THz photons which can be detected using an SPD, allowing ATIs to detect dark matter. We showed that antiferromagnetic Fe-doped Bi_2Se_3 satisfies the three Wilczek criteria described earlier, and can be used to realize a DA detector in the 0.7 to 3.5 meV range. Figure 4 shows the projected reach of three possible schemes with different effective volumes. Varying the applied B field scans the resonant frequency, giving sensitivity to axion dark matter in a parameter space inaccessible to other methods. Future work on the material characteristics (such as the anisotropy field strength) can allow for a wider range of DA mass detection.

We acknowledge useful discussions with Francesca Day, Joerg Jaeckel, Alexander Millar, Jens Niemeyer, David Tanner, Chris Weber, William Witzak-Krempa, and Ariel Zhitnisky. M. A. and D. J. E. M. are supported by the Alexander von Humboldt Foundation and the German Federal Ministry of Education and Research. K. C. F. acknowledges support from Army Research Office under Cooperative Agreement No. W911NF-17-1-0574. L. Š. acknowledges support from the EU FET Open RIA Grant No. 766566.

* david.marsh@uni-goettingen.de

- [1] G. Bertone, D. Hooper, and J. Silk, *Phys. Rep.* **405**, 279 (2005).
- [2] Planck Collaboration *et al.*, *Astron. Astrophys.* **594**, A13 (2016).
- [3] F. Iocco, M. Pato, and G. Bertone, *Nat. Phys.* **11**, 245 (2015).
- [4] K. Olive *et al.* (Particle Data Group), *Chin. Phys. C* **38**, 090001 (2014).
- [5] L. F. Abbott and P. Sikivie, *Phys. Lett.* **120B**, 133 (1983).
- [6] M. Dine and W. Fischler, *Phys. Lett.* **120B**, 137 (1983).
- [7] J. Preskill, M. B. Wise, and F. Wilczek, *Phys. Lett.* **120B**, 127 (1983).

- [8] J. E. Kim and G. Carosi, *Rev. Mod. Phys.* **82**, 557 (2010).
- [9] A. Ringwald, *Phys. Dark Universe* **1**, 116 (2012).
- [10] P. W. Graham, I. G. Irastorza, S. K. Lamoreaux, A. Lindner, and K. A. van Bibber, *Annu. Rev. Nucl. Part. Sci.* **65**, 485 (2015).
- [11] D. J. E. Marsh, *Phys. Rep.* **643**, 1 (2016).
- [12] R. D. Peccei and H. R. Quinn, *Phys. Rev. Lett.* **38**, 1440 (1977).
- [13] S. Weinberg, *Phys. Rev. Lett.* **40**, 223 (1978).
- [14] F. Wilczek, *Phys. Rev. Lett.* **40**, 279 (1978).
- [15] J. M. Pendlebury, S. Afach, N. J. Ayres, C. A. Baker, G. Ban, G. Bison, K. Bodek, M. Burghoff, P. Geltenbort, K. Green *et al.*, *Phys. Rev. D* **92**, 092003 (2015).
- [16] S. Borsanyi *et al.*, *Nature (London)* **539**, 69 (2016).
- [17] S. J. Asztalos *et al.*, *Phys. Rev. Lett.* **104**, 041301 (2010).
- [18] N. Du, N. Force, R. Khatiwada, E. Lentz, R. Ottens, L. J. Rosenberg, G. Rybka, G. Carosi, N. Woollett, D. Bowring *et al.* (ADMX Collaboration), *Phys. Rev. Lett.* **120**, 151301 (2018).
- [19] J. E. Kim, *Phys. Rev. Lett.* **43**, 103 (1979).
- [20] M. A. Shifman, A. I. Vainshtein, and V. I. Zakharov, *Nucl. Phys.* **B166**, 493 (1980).
- [21] A. Zhitnitsky, *Sov. J. Nucl. Phys.* **31**, 260 (1980).
- [22] M. Dine, W. Fischler, and M. Srednicki, *Phys. Lett. B* **104**, 199 (1981).
- [23] G. G. Raffelt, Astrophysical axion bounds, in *Axions*, edited by M. Kuster, G. Raffelt, and B. Beltrán, Lecture Notes in Physics, Vol. 741 (Springer Verlag, Berlin, 2008), p. 51.
- [24] A. Ayala, I. Domínguez, M. Giannotti, A. Mirizzi, and O. Straniero, *Phys. Rev. Lett.* **113**, 191302 (2014).
- [25] M. Kawasaki, K. Saikawa, and T. Sekiguchi, *Phys. Rev. D* **91**, 065014 (2015).
- [26] A. H. Córscico *et al.*, *Mon. Not. R. Astron. Soc.* **424**, 2792 (2012).
- [27] M. Arik *et al.*, *Phys. Rev. Lett.* **112**, 091302 (2014).
- [28] K. Ehret *et al.*, *Phys. Lett. B* **689**, 149 (2010).
- [29] A. Arvanitaki and A. A. Geraci, *Phys. Rev. Lett.* **113**, 161801 (2014).
- [30] Y. J. Kim, P.-H. Chu, and I. Savukov, *Phys. Rev. Lett.* **121**, 091802 (2018).
- [31] L. Zhong, S. Al Kenany, K. M. Backes, B. M. Brubaker, S. B. Cahn, G. Carosi, Y. V. Gurevich, W. F. Kindel, S. K. Lamoreaux, K. W. Lehnert *et al.* (HAYSTAC Collaboration), *Phys. Rev. D* **97**, 092001 (2018).
- [32] B. T. McAllister, G. Flower, E. N. Ivanov, M. Goryachev, J. Bourhill, and M. E. Tobar, *Phys. Dark Universe* **18**, 67 (2017).
- [33] N. Crescini *et al.*, *Eur. Phys. J. C* **78**, 703 (2018); **78**, 813(E) (2018).
- [34] A. Caldwell, Gia Dvali, Béla Majorovits, Alexander Millar, Georg Raffelt, Javier Redondo, Olaf Reimann, Frank Simon, and Frank Steffen (MADMAX Working Group), *Phys. Rev. Lett.* **118**, 091801 (2017).
- [35] P. Sikivie, N. Sullivan, and D. B. Tanner, *Phys. Rev. Lett.* **112**, 131301 (2014).
- [36] Y. Kahn, B. R. Safdi, and J. Thaler, *Phys. Rev. Lett.* **117**, 141801 (2016).
- [37] B. T. McAllister, G. Flower, E. N. Ivanov, M. Goryachev, J. Bourhill, and M. E. Tobar, *Phys. Dark Universe* **18**, 67 (2017).
- [38] R. Barbieri, C. Braggio, G. Carugno, C. S. Gallo, A. Lombardi, A. Ortolan, R. Pengo, G. Ruoso, and C. C. Speake, *Phys. Dark Universe* **15**, 135 (2017).
- [39] D. Budker, P. W. Graham, M. Ledbetter, S. Rajendran, and A. O. Sushkov, *Phys. Rev. X* **4**, 021030 (2014).
- [40] C. Abel, N. J. Ayres, G. Ban, G. Bison, K. Bodek, V. Bondar, M. Daum, M. Fairbairn, V. V. Flambaum, P. Geltenbort *et al.*, *Phys. Rev. X* **7**, 041034 (2017).
- [41] Y. V. Stadnik and V. V. Flambaum, *Phys. Rev. D* **89**, 043522 (2014).
- [42] S. L. Adler, *Phys. Rev.* **177**, 2426 (1969).
- [43] J. S. Bell and R. Jackiw, *Nuovo Cimento A* **60**, 47 (1969).
- [44] A. J. Millar, G. G. Raffelt, J. Redondo, and F. D. Steffen, *J. Cosmol. Astropart. Phys.* **01** (2017) 061.
- [45] R. Li, J. Wang, X.-L. Qi, and S.-C. Zhang, *Nat. Phys.* **6**, 284 (2010).
- [46] A. M. Essin, J. E. Moore, and D. Vanderbilt, *Phys. Rev. Lett.* **102**, 146805 (2009).
- [47] A. Malashevich, S. Coh, I. Souza, and D. Vanderbilt, *Phys. Rev. B* **86**, 094430 (2012).
- [48] J. Wang, R. Li, S.-C. Zhang, and X.-L. Qi, *Phys. Rev. Lett.* **106**, 126403 (2011).
- [49] X. Wan, A. Vishwanath, and S. Y. Savrasov, *Phys. Rev. Lett.* **108**, 146601 (2012).
- [50] J. Wang, B. Lian, and S.-C. Zhang, *Phys. Rev. B* **93**, 045115 (2016).
- [51] L. Wu, M. Salehi, N. Koirala, J. Moon, S. Oh, and N. P. Armitage, *Science* **354**, 1124 (2016).
- [52] V. Dziom *et al.*, *Nat. Commun.* **8**, 15197 (2017).
- [53] K. N. Okada, Y. Takahashi, M. Mogi, R. Yoshimi, A. Tsukazaki, K. S. Takahashi, N. Ogawa, M. Kawasaki, and Y. Tokura, *Nat. Commun.* **7**, 12245 (2016).
- [54] M. Mogi, M. Kawamura, R. Yoshimi, A. Tsukazaki, Y. Kozuka, N. Shirakawa, K. S. Takahashi, M. Kawasaki, and Y. Tokura, *Nat. Mater.* **16**, 516 (2017).
- [55] D. Xiao *et al.*, *Phys. Rev. Lett.* **120**, 056801 (2018).
- [56] S. Grauer, K. M. Fijalkowski, S. Schreyeck, M. Winnerlein, K. Brunner, R. Thomale, C. Gould, and L. W. Molenkamp, *Phys. Rev. Lett.* **118**, 246801 (2017).
- [57] A. Sekine and K. Nomura, *Phys. Rev. Lett.* **116**, 096401 (2016).
- [58] H. Terças, J. D. Rodrigues, and J. T. Mendonça, *Phys. Rev. Lett.* **120**, 181803 (2018).
- [59] M. Srednicki, *Nucl. Phys.* **B260**, 689 (1985).
- [60] P. W. Graham and S. Rajendran, *Phys. Rev. D* **88**, 035023 (2013).
- [61] Y. Hochberg, Y. Kahn, M. Lisanti, K. M. Zurek, A. G. Grushin, R. Ilan, S. M. Griffin, Z.-F. Liu, S. F. Weber, and J. B. Neaton, *Phys. Rev. D* **97**, 015004 (2018).
- [62] F. Wilczek, *Phys. Rev. Lett.* **58**, 1799 (1987).
- [63] M. Z. Hasan and C. L. Kane, *Rev. Mod. Phys.* **82**, 3045 (2010).
- [64] L. Šmejkal, J. Železný, J. Sinova, and T. Jungwirth, *Phys. Rev. Lett.* **118**, 106402 (2017).
- [65] L. Šmejkal, Y. Mokrousov, B. Yan, and A. H. MacDonald, *Nat. Phys.* **14**, 242 (2018).
- [66] A. Sekine and K. Nomura, *J. Phys. Soc. Jpn.* **83**, 104709 (2014).
- [67] L. Fu and C. L. Kane, *Phys. Rev. Lett.* **100**, 096407 (2008).
- [68] S. Rachel, *Rep. Prog. Phys.* **81**, 116501 (2018).

- [69] C. Kittel, *Phys. Rev.* **82**, 565 (1951).
- [70] P. L. Richards, *J. Appl. Phys.* **35**, 850 (1964).
- [71] F. Keffer, H. Kaplan, and Y. Yafet, *Am. J. Phys.* **21**, 250 (1953).
- [72] H.-J. Kim *et al.*, *Phys. Rev. Lett.* **110**, 136601 (2013).
- [73] J.-M. Zhang, W. Zhu, Y. Zhang, D. Xiao, and Y. Yao, *Phys. Rev. Lett.* **109**, 266405 (2012).
- [74] J.-M. Zhang, W. Ming, Z. Huang, G.-B. Liu, X. Kou, Y. Fan, K. L. Wang, and Y. Yao, *Phys. Rev. B* **88**, 235131 (2013).
- [75] A. J. Millar, J. Redondo, and F. D. Steffen, *J. Cosmol. Astropart. Phys.* **10** (2017) 006; **05** (2018) E02.
- [76] K. Van Bibber, N. R. Dagdeviren, S. E. Koonin, A. K. Kerman, and H. N. Nelson, *Phys. Rev. Lett.* **59**, 759 (1987).
- [77] M. Z. Hasan and C. L. Kane, *Rev. Mod. Phys.* **82**, 3045 (2010).
- [78] D. Horns, J. Jaeckel, A. Lindner, A. Lobanov, J. Redondo, and A. Ringwald, *J. Cosmol. Astropart. Phys.* **04** (2013) 016.
- [79] P. Sikivie, *Phys. Rev. Lett.* **51**, 1415 (1983).
- [80] A. J. Sievers and M. Tinkham, *Phys. Rev.* **129**, 1566 (1963).
- [81] S. P. Bayrakci, D. A. Tennant, Ph. Leininger, T. Keller, M. C. R. Gibson, S. D. Wilson, R. J. Birgeneau, and B. Keimer, *Phys. Rev. Lett.* **111**, 017204 (2013).
- [82] B. Náfrádi, T. Keller, H. Manaka, U. Stuhr, A. Zheludev, and B. Keimer, *Phys. Rev. B* **87**, 020408 (2013).
- [83] A. Zheludev, V. O. Garlea, L.-P. Regnault, H. Manaka, A. Tselik, and J.-H. Chung, *Phys. Rev. Lett.* **100**, 157204 (2008).
- [84] A. Little *et al.*, *Phys. Rev. Lett.* **119**, 227201 (2017).
- [85] A. B. Harris, D. Kumar, B. I. Halperin, and P. C. Hohenberg, *Phys. Rev. B* **3**, 961 (1971).
- [86] S. M. Rezende, R. L. Rodríguez-Suárez, and A. Azevedo, *Phys. Rev. B* **93**, 014425 (2016).
- [87] Y. Shiomi, R. Takashima, D. Okuyama, G. Gitgeatpong, P. Piyawongwatthana, K. Matan, T. J. Sato, and E. Saitoh, *Phys. Rev. B* **96**, 180414(R) (2017).
- [88] S. K. Lamoreaux, K. A. van Bibber, K. W. Lehnert, and G. Carosi, *Phys. Rev. D* **88**, 035020 (2013).
- [89] S. Komiyama, O. Astafiev, V. Antonov, T. Kutsuwa, and H. Hirai, *Nature (London)* **403**, 405 (2000).
- [90] E. D. Walsh, D. K. Efetov, G.-H. Lee, M. Heuck, J. Crossno, T. A. Ohki, P. Kim, D. Englund, and K. C. Fong, *Phys. Rev. Applied* **8**, 024022 (2017).
- [91] C. A. J. O'Hare and A. M. Green, *Phys. Rev. D* **95**, 063017 (2017).
- [92] H. Peng *et al.*, *Nucl. Instrum. Methods Phys. Res., Sect. A* **444**, 569 (2000).
- [93] P. Losták, Č. Drašar, I. Klichová, J. Navrátil, and T. Černohorský, *Phys. Status Solidi B* **200**, 289 (1997).
- [94] X. Wei, J. Zhang, B. Zhao, Y. Zhu, and Z. Yang, *Phys. Lett. A* **379**, 417 (2015).
- [95] R. Bradley, J. Clarke, D. Kinion, L. J. Rosenberg, K. van Bibber, S. Matsuki, M. Mück, and P. Sikivie, *Rev. Mod. Phys.* **75**, 777 (2003).

Correction: The Cooperative Agreement number contained an error and has been fixed.