

ALMA Survey of Orion Planck Galactic Cold Clumps (ALMASOP) I. Detection of New Hot Corinos with ACA

SHIH-YING HSU,^{1,2} SHENG-YUAN LIU,² TIE LIU,³ DIPEN SAHU,² NAOMI HIRANO,² CHIN-FEI LEE,² KEN'ICHI TATEMATSU,^{4,5}
GWANJEONG KIM,⁴ MIKA JUVELA,⁶ PATRICIO SANHUEZA,⁷ JINHUA HE,^{8,9,10} DOUG JOHNSTONE,^{11,12} SHENG-LI QIN,¹³
LEONARDO BRONFMAN,¹⁰ HUEI-RU VIVIEN CHEN,¹⁴ SOMNATH DUTTA,² DAVID J. EDEN,¹⁵ KAI-SYUN JHAN,^{1,2} KEE-TAE KIM,^{16,17}
YI-JEHNG KUAN,^{18,2} WOJIN KWON,^{19,16} CHANG WON LEE,^{16,17} JEONG-EUN LEE,²⁰ ANTHONY MORAGHAN,² M. G. RAWLINGS,²¹
HSIEN SHANG,² ARCHANA SOAM,²² M.A. THOMPSON,²³ ALESSIO TRAFICANTE,²⁴ YUEFANG WU,^{25,26} YAO-LUN YANG,²⁷ AND
QIZHOU ZHANG²⁸

¹National Taiwan University (NTU), Taiwan (R.O.C.)

²Institute of Astronomy and Astrophysics, Academic Sinica, Taipei, Taiwan (R.O.C.)

³Key Laboratory for Research in Galaxies and Cosmology, Shanghai Astronomical Observatory, Chinese Academy of Sciences, 80 Nandan Road, Shanghai 200030, Peoples Republic of China

⁴Nobeyama Radio Observatory, National Astronomical Observatory of Japan, National Institutes of Natural Sciences, 462-2 Nobeyama, Minamimaki, Minamisaku, Nagano 384-1305, Japan

⁵Department of Astronomical Science, SOKENDAI (The Graduate University for Advanced Studies), 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan

⁶Department of Physics, P.O.Box 64, FI-00014, University of Helsinki, Finland

⁷National Astronomical Observatory of Japan, National Institutes of Natural Sciences, 2-21-1 Osawa, Mitaka, Tokyo 181-8588, Japan

⁸Yunnan Observatories, Chinese Academy of Sciences, 396 Yangfangwang, Guandu District, Kunming, 650216, P. R. China

⁹Chinese Academy of Sciences South America Center for Astronomy, National Astronomical Observatories, CAS, Beijing 100101, China

¹⁰Departamento de Astronomía, Universidad de Chile, Casilla 36-D, Santiago, Chile

¹¹NRC Herzberg Astronomy and Astrophysics, 5071 West Saanich Rd, Victoria, BC, V9E 2E7, Canada

¹²Department of Physics and Astronomy, University of Victoria, Victoria, BC, V8P 5C2, Canada

¹³Department of Astronomy, Yunnan University, and Key Laboratory of Astroparticle Physics of Yunnan Province, Kunming, 650091, Peoples Republic of China

¹⁴Institute of Astronomy and Department of Physics, National Tsing Hua University, Hsinchu 30013, Taiwan

¹⁵Astrophysics Research Institute, Liverpool John Moores University, iC2, Liverpool Science Park, 146 Brownlow Hill, Liverpool, L3 5RF, UK.

¹⁶Korea Astronomy and Space Science Institute (KASI), 776 Daedeokdae-ro, Yuseong-gu, Daejeon 34055, Republic of Korea

¹⁷University of Science and Technology, Korea (UST), 217 Gajeong-ro, Yuseong-gu, Daejeon 34113, Republic of Korea

¹⁸Department of Earth Sciences, National Taiwan Normal University, Taipei, Taiwan (R.O.C.)

¹⁹Department of Earth Science Education, Seoul National University (SNU), 1 Gwanak-ro, Gwanak-gu, Seoul 08826, Republic of Korea

²⁰School of Space Research, Kyung Hee University, 1732, Deogyong-Daero, Giheung-gu Yongin-shi, Gyeonggi-do 17104, Korea

²¹East Asian Observatory, 660 N. Aohk Place, University Park, Hilo, HI 96720, USA

²²SOFIA Science Center, USRA, NASA Ames Research Center, M.S.-12, N232, Moffett Field, CA 94035, USA

²³Centre for Astrophysics Research, School of Physics Astronomy & Mathematics, University of Hertfordshire, College Lane, Hatfield, AL10 9AB, UK.

²⁴INAF-IAPS, via Fosso del Cavaliere, 100. 00133, Rome, IT

²⁵Department of Astronomy, School of Physics, Peking University, Beijing, 1000871, PR China

²⁶Kavli Institute for Astronomy and Astrophysics, Peking University, Beijing, 100871, PR China

²⁷Department of Astronomy, University of Virginia, Charlottesville, VA 22904, USA

²⁸Center for Astrophysics — Harvard & Smithsonian, 60 Garden Street, Cambridge, MA 02138, USA

(Accepted June 20, 2020)

Submitted to ApJ

ABSTRACT

We report the detection of four new hot corino sources, G211.47-19.27S, G208.68-19.20N1, G210.49-19.79W and G192.12-11.10 from a survey study of Planck Galactic Cold Clumps in the Orion Molecular Cloud Complex with the Atacama Compact Array (ACA). Three sources had been identified as low mass Class 0 protostars in the Herschel Orion Protostar Survey (HOPS). One source in the λ Orionis region is firstly reported as a

protostellar core. We have observed abundant complex organic molecules (COMs), primarily methanol but also other oxygen-bearing COMs (in G211.47-19.27S and G208.68-19.20N1) and the molecule of prebiotic interest NH_2CHO (in G211.47-19.27S), signifying the presence of hot corinos. While our spatial resolution is not sufficient for resolving most of the molecular emission structure, the large linewidth and high rotational temperature of COMs suggest that they likely reside in the hotter and innermost region immediately surrounding the protostar. In G211.47-19.27S, the D/H ratio of methanol ($[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$) and the $^{12}\text{C}/^{13}\text{C}$ ratio of methanol ($[\text{CH}_3\text{OH}]/[^{13}\text{CH}_3\text{OH}]$) are comparable to those of other hot corinos. Hydrocarbons and long carbon-chain molecules such as $\text{c-C}_3\text{H}_2$ and HCCCN are also detected in the four sources, likely tracing the outer and cooler molecular envelopes.

Keywords: astrochemistry — ISM: molecules — stars: formation and low-mass

1. INTRODUCTION

Many Class 0/I low- (and also intermediate-) mass protostellar cores show considerable chemical diversity. Those cores characterized by the presence of abundant saturated complex organic molecules (COMs) within a warm (~ 100 K) and compact (< 100 au) region around the central protostar are called “hot corinos” (Ceccarelli 2004). Hot corinos seem qualitatively similar to “hot cores” also characterized by abundant COMs but associated with larger (~ 1000 au) and warmer (~ 300 K) regions in high-mass star formation sites; however, the abundances of the COMs in hot corinos can be, in some cases, orders of magnitude higher (fractional abundance respect to the hydrogen molecule $X \sim 10^{-7} - 10^{-9}$) than those in hot cores (Ospina-Zamudio et al. 2018).

The current hot corino formation scenario can be described by three phases (Herbst & van Dishoeck 2009). Initially, atoms and molecules in the gas phase of prestellar cores accrete onto dust grains and the 0th generation COMs (e.g. the COM precursor H_2CO , methanol CH_3OH , and possibly other COMs) are produced by grain-surface chemistry during this “cold phase.” In the “warm-up phase”, the photodissociated radicals (e.g. methoxy radical CH_3O and formyl radical HCO of CH_3OH and H_2CO , respectively) form larger complex molecules (e.g. methyl formate HCOOCH_3 and formic acid HCOOH) as the first-generation molecules via radical-radical reactions in the ice mantles. The “hot corino phase” takes place when the temperature reaches ~ 100 K from 10 K in the cold phase. The icy mantles completely sublimate into the gas and the second-generation molecules are consequently produced via gas-phase reactions.

So far a good number of hot corinos have been identified, including IRAS 16293-2422 in the Rho Ophiuchi cloud complex (Cazaux et al. 2003), IRAS4A2, IRAS2A, IRAS4B and SVS13-A in the Perseus molecular cloud (Bottinelli et al. 2004; Jorgensen et al. 2005; López-Sepulcre et al. 2017; Bottinelli et al. 2007; Bianchi et al. 2018), B335 in Lynd L663 (Imai et al. 2016), B1b in Barnard 1 (Lefloch et al. 2018), Cep E-mm in the Cepheus E molecular cloud (Ospina-Zamudio et al. 2018), L483 in Serpens-Aquila

Rift (Oya et al. 2017; Jacobsen et al. 2019), Ser-emb 1, Ser-emb 8 and Ser-emb 17 in the Serpens Cluster B (Martín-Doménech et al. 2019; Bergner et al. 2019), HH-212 in the Orion Molecular Cloud Complex (Codella et al. 2016; Lee et al. 2017) and BHR-71 IRS1 in the BHR-71 Bok globule (Yang et al. 2020). In addition, a hot-corino-like atmosphere was implied toward IRAS4A1 in the Perseus 103 molecular cloud (Sahu et al. 2019). In particular, the hot corino nature of most of these corinos were found through case studies toward specific sources rather than from surveys.

In addition to the hot corino chemistry, there is also the so-called “warm-carbon-chain chemistry (WCCC)” associated with protostellar warm (~ 30 K) infalling envelopes being abundant in unsaturated carbon-chain molecules (Sakai et al. 2008, 2009). Although the hot corino chemistry and the warm-carbon chain chemistry exhibit two distinct behaviors, there are sources bearing signatures of both COMs and long-carbon chain molecules. Higuchi et al. (2018) carried out a survey toward 36 Class 0/I protostars in the Perseus molecular cloud complex and showed that the abundance ratios between CH_3OH , a hot corino tracer, and ethynyl (C_2H) and cyclopropenylidene ($\text{c-C}_3\text{H}_2$), both being proxies of WCCC, range within a continuous spectrum of 1–2 orders of magnitude. While there is no distinct separation between the two types of chemical signatures, the column density ratios between ethynyl and methanol ($[\text{C}_2\text{H}]/[\text{CH}_3\text{OH}]$) appear to be correlated with the core locations in the cloud complex (Higuchi et al. 2018).

The Planck Galactic Cold Clumps (PGCC) catalog is an all-sky catalog consisting of cold clump candidates (Planck Collaboration et al. 2016). At an angular scale of $\sim 5'$, the Planck’s resolution, these PGCCs appear dense (with molecular hydrogen column density $N(\text{H}_2) > 10^{20} \text{ cm}^{-2}$) and cold (10–20 K) regions potentially harboring star formation at their very early stages. Survey observations in $850 \mu\text{m}$ were previously conducted with the Submillimetre Common User Bolometer Array-2 (SCUBA-2) at the 15m James Clerk Maxwell Telescope (JCMT) toward 96 dense PGCCs (clump-averaged column density larger than $5 \times 10^{20} \text{ cm}^{-2}$) in Orion A, B and λ Orionis (Yi et al. 2018), and as part

of the JCMT large program ‘‘SCOPE: SCUBA-2 Continuum Observations of Pre-protostellar Evolution (Liu et al. 2018; Eden et al. 2019). This 850 μm (dust) continuum survey identified 119 protostellar and starless cores within 40 Orion PGCCs (Yi et al. 2018). These cores were further observed with the Nobeyama Radio Observatory (NRO) 45m telescope for studying their evolutionary stages gauged by, for example, their N_2D^+ abundances (Tatematsu et al. 2017, Kim et al. 2020, submitted). On the basis of these surveys, we selected 72 dense and compact 850 μm continuum cores at early stages (starless cores with intense N_2D^+ emission and Class 0 protostellar cores with and without intense N_2D^+ emission) and carried out observations with the Atacama Compact Array (ACA) within an Atacama Large Millimeter/submillimeter Array (ALMA) project. This project, ALMA Survey of Orion Planck Galactic Cold Clumps (ALMASOP), aims to probe the onset of star formation (See Sect. 2). We conducted chemical studies on the sample and report in this paper the finding of four sources, being rich in molecular lines and harboring saturated COMs. Information about the ALMA observations over the full sample will be reported in a separate paper (Dutta et al. in preparation).

2. OBSERVATIONS

The observations were carried out with the ACA, which is an interferometer composed of twelve 7-meter antennas. They were conducted as a part of the ALMA Cycle 6 project, (#2018.1.00302.S; PI: Tie Liu), aiming to study the fragmentation of dense cores in the Orion Molecular Cloud Complex.

The longest baseline was 49 m, which was about 37.7 k λ . The HPBW of the synthesized beam was $\sim 5.6''$ and the field of view was $\sim 43.2''$. Four spectral windows, centered at 216.6, 218.9, 231.0 and 233.0 GHz with an uniform bandwidth of 1875 MHz ($\sim 2500 \text{ km s}^{-1}$) and a resolution of $\sim 1.129 \text{ MHz}$ ($\sim 1.5 \text{ km s}^{-1}$) were set.

The data were calibrated with the Common Astronomy Software Applications package (CASA, McMullin 2007) version 5.4.0-68 and its pipeline version 42030M (Pipeline-CASA54-P1-B). In the pipeline, the images were processed through tclean with auto-masking and the robust parameter of the Briggs weighting set to 0.5. See ALMA Science Pipeline website¹ for other parameters of auto-masking (e.g. sidelobethreshold=1.25, noisethreshold=5.0 and negativethreshold=0.0).

We report the result of four sources, which we identified as hot corinos within the whole sample (See Sect. 4.2). They are G211.47-19.27S, G208.68-19.20N1, G210.49-19.79W and G192.12-11.10, hereafter G211, G208, G210 and G192, respectively. See Table 1 for the coordinates, date, on-source integration time, calibrators of the observations, the sizes and

position angles of beams, and the root-mean-square (RMS) noise of the continuum and line images.

We note that in this paper we use only the 7m ACA data for the hot corino identification. The ALMA 12m array data, which provide better angular resolutions, will be employed for further analyses, such as the structures of the line emitting regions, in a forthcoming paper.

3. RESULTS

3.1. Dust Continuum

The continuum images of the four sources are shown in Figure 1. All four sources are well detected with a single dominant components in the map centers and their peak brightness temperatures are all less than 1 K. Such brightness temperatures are much less than the expected physical temperature (greater than 10 K in protostars) so their continuum emission is optically thin. The images for the two brighter targets (G211 and G208) are severely dynamical-range limited (i.e. the residual sidelobe features dominate over the ideal theoretical noise, $5\times$ to $10\times$ in our cases), hence noisier than maps for the other two targets as indicated in Table 1. In addition to the compact component, there exists also an extended component in G210 and it is possibly the remnant of the envelope. We applied a 2D Gaussian fit in CASA to the continuum and based on the fitting result, including the source-averaged flux, the peak positions, the integrated flux density and the apparent and de-convolved angular sizes, we calculated the molecular hydrogen column density (N_{H_2}), the mass (M) and the radius (R) of the sources with formulae modified from Eq. A.26 and A.30 in Kauffmann et al. (2008) for optically thin continuum emission.

The source-averaged molecular hydrogen column density N_{H_2} can be derived via:

$$N_{\text{H}_2} = \frac{F_{\text{CONT}}^{\text{Int}}}{\Omega_S \mu_{\text{H}_2} m_{\text{H}} \kappa_{\nu} B_{\nu}(T_{\text{d}})} \quad (1)$$

where $F_{\text{CONT}}^{\text{Int}}$ is the total flux, Ω_S is the size of the deconvolved source image, μ_{H_2} is the molecular weight per hydrogen molecule ~ 2.8 , m_{H} is the mass of the hydrogen atom, $B_{\nu}(T_{\text{d}})$ is the black-body radiation function evaluated at the dust temperature T_{d} . κ_{ν} is the dust opacity in the form of $\kappa_{\nu} = 0.1(\nu/1 \text{ THz})^{\beta} \text{ cm}^2 \text{ g}^{-1}$ where β is the dust opacity index (Beckwith et al. 1990). The indices β were assumed to be 1.70, which is the typical opacity index of cold clumps in the submillimeter band (Juvella et al. 2018) and the resulting κ_{ν} at 1.3 mm is $0.0083 \text{ cm}^2 \text{ g}^{-1}$. Since the continuum emission of all the four sources are marginally resolved (i.e. the observed source size is larger but comparable to the synthesized beam size), the derived column densities could be lower limits.

The mass is estimated via:

$$M = \frac{F_{\nu} D^2}{\kappa_{\nu} B_{\nu}(T_{\text{d}})} \quad (2)$$

¹ <https://almascience.eso.org/processing/science-pipeline>

Table 1. The parameters of the observations.

Source Name	α_{J2000}	δ_{J2000}	Total On-source Integration Time	Beam $\theta_{\max}, \theta_{\min}, \text{PA}$	$\sigma_{\text{CONT}}, \sigma_{\text{Chn}}$ mJy beam ⁻¹ (mK)	Date	Calibrator
							Bandpass & Flux, Phase
G211.47-19.27S	05:39:56.097	-07:30:28.403	210 sec	8.79'', 4.20'', -76.5°	8.6 (5.4), 37.1 (23.3)	11/19	J0522-3627, J0607-0834
							11/21 J0854+2006, J0607-0834
G208.68-19.20N1	05:35:23.486	-05:01:31.583	500 sec	7.60'', 4.00'', -83.3°	13.5 (10.3), 27.5 (20.9)	11/27	J0423-0120, J0542-0913
G210.49-19.79W	05:36:18.860	-06:45:28.035	300 sec	7.65'', 4.02'', -82.4°	1.9 (1.5), 27.3 (20.5)	11/27	J0854+2006, J0607-0834
							11/27 J0522-3627, J0542-0913
							11/28 J0423-0120, J0542-0913
G192.12-11.10	05:32:19.540	+12:49:40.190	480 sec	6.86'', 4.81'', 80.7°	1.2 (0.9), 20.1 (14.1)	11/21	J0423-0120, J0530+1331
							11/26 J0423-0120, J0530+1331

NOTE— All the measurements were made in 2018. The observations toward G208 and G210 were separated into five executions and G192 observation toward were separated into two executions. The θ_{\max} and θ_{\min} are the FWHM of the synthesized beam along the major and minor axes, respectively. The PA is the beam position angle. The σ_{CONT} and σ_{Chn} are the root-mean-square (RMS) noises of the measurement spanning over 7.5 GHz (continuum) and 1.129 MHz (spectral resolution), respectively. The brightness temperatures of σ_{CONT} and σ_{Chn} are derived with 230 GHz and the synthesized beam area of each source (see Table 2).

where D is the distance to the source adopted from the PGCC catalog (Planck Collaboration et al. 2016). The sources are marginally resolved and the radius of the sources were estimated via

$$R = \frac{1}{2} D \theta_S \quad (3)$$

where θ_S is the deconvolved FWHM size of the source derived from the 2D fit to the continuum data. All the derived quantities are presented in Table 2. We note that the gas column density and mass are estimated with a range of dust temperatures at 50K, 100K, and 150K.

According to the coordinates of the continuum peaks, we searched for known young stellar objects (YSOs) within 5'' using SIMBAD² (Wenger et al. 2000). We found that three of the four sources are likely to be associated with the Herschel Orion Protostar Survey (HOPS) objects (Manoj et al. 2013), HOPS 288 for G211, HOPS 87 for G208 and HOPS 168 for G210, which are all located in the Orion A molecular cloud (see Table 3). First, the offsets between each source and its corresponding HOPS object are small (0.7'', 2.17'' and 0.84'' for G211, G208 and G210, respectively).

Second, there are no other compact sources within each of their fields of view. In contrast, G192 in λ Orionis is firstly identified as a YSO. See Sect. 4.1 for more discussions.

3.2. Molecular Emission

We extracted the spectra at the continuum peak in the four spectral datacubes for each of the 72 sources in the project (see Sect. 1). We then scanned all the samples and identified only those containing rich spectral features including

particularly multiple CH₃OH lines with excitation temperature higher than 100 K for further analysis. With the above criteria, we identified four sources (G211, G208, G210, and G192) and present their spectra, respectively, in Fig. A1, A2, A3, and A4. In total, around 154 molecular lines in G211, 80 lines in G208, 38 lines in G210 and 30 lines in G192 above $3\sigma_{\text{Chn}}$ were recognized. We further applied Gaussian fitting over those transitions to obtain the observed frequency (f^{obs}), peak brightness temperature (T_p), velocity width (Δv), and the integrated intensity (w) of the identified lines. The result of the Gaussian fitting to each line of each source is in Table B1, B2, B3 and B4 for G211, G208, G210, and G192, respectively.

For line identifications, we searched the Splatalogue³, which is an online database including the Jet Propulsion Laboratory Molecular Spectroscopy (JPL⁴, Pickett et al. 1998) and the Cologne Database for Molecular Spectroscopy (CDMS⁵, Miller et al. 2005) catalog for astronomical spectroscopy, for possible molecule candidates. We detected 25 molecular species (including isotopologues) in G211, 18 in G208, 18 in G210 and 16 in G192. Some species were detected in all the four sources including CO, C¹⁸O, OCS, ¹³CS, H₂S, HCCCN, DCN, DCO⁺, D₂CO, H₂CO, c-C₃H₂, CCD, ³⁴SO, HCOOH and CH₃OH. Two molecules were detected in three of the four sources. They are HNCO (not in G208) and SiO (not in G192). CH₃CHO and HCOOCH₃ were detected in G211 and G208, but were not found in

³ <https://www.cv.nrao.edu/php/splat/>

⁴ <https://spec.jpl.nasa.gov/>

⁵ <https://cdms.astro.uni-koeln.de/>

² <http://simbad.u-strasbg.fr/simbad/>

Table 2. The parameters and result of the continuum analysis.

Source Name	θ_A ($''$)	θ_S ($''$)	$F_{\text{CONT}}^{\text{Int}}$ (mJy)	D (pc)	N_{H_2} (cm^{-2})			M (M_{\odot})			R (au)
					50 K	100 K	150 K	50 K	100 K	150 K	
G211.47-19.27S	6.07	2.22	552	415	1.63e+24	7.73e+23	5.06e+23	0.82	0.39	0.25	461
G208.68-19.20N1	5.51	2.51	1530	415	3.55e+24	1.68e+24	1.10e+24	2.28	1.08	0.71	520
G210.49-19.79W	5.55	3.42	185	415	2.31e+23	1.09e+23	7.16e+22	0.28	0.13	0.09	709
G192.12-11.10	5.74	2.87	199	400	3.52e+23	1.66e+23	1.09e+23	0.28	0.13	0.09	575

NOTE— θ_A is the FWHM of the synthesized beam. θ_S is the FWHM of the deconvolved source image. θ_A and θ_S are the geometric mean of their maximum and minimum FWHM values, i.e. $\theta = \sqrt{\theta_{\text{max}}\theta_{\text{min}}}$. $F_{\text{CONT}}^{\text{Int}}$ is the total flux integrated over the source. N_{H_2} , M and R are the calculated column density of molecular hydrogen, the mass, and the radius of the sources, respectively. The N_{H_2} and M are estimated at three different dust temperatures, 50 K, 100 K and 150 K. Note that column densities listed here could be lower limits. See Sect. 3.1

Table 3. The coordinates of the continuum peaks and the information of their corresponding HOPS objects and the protostellar classifications.

Source Name	Cloud	$\alpha_{\text{J2000}}^{\text{peak}}$	$\delta_{\text{J2000}}^{\text{peak}}$	HOPS (Furlan et al. 2016)					
				Index	α_{J2000}	δ_{J2000}	Class	L_{bol} (L_{\odot})	T_{bol} (K)
G211.47-19.27S	Orion A	05:39:55.988	-07:30:27.593	HOPS 288	05:39:55.95	-07:30:28.0	0	135.47	48.6
G208.68-19.20N1	Orion A	05:35:23.434	-05:01:30.803	HOPS 87	05:35:23.47	-05:01:28.7	0	36.49	38.1
G210.49-19.79W	Orion A	05:36:18.965	-06:45:23.355	HOPS 168	05:36:18.93	-06:45:22.7	0	48.07	54.0
G192.12-11.10	λ Orionis	05:32:19.345	+12:49:41.140

the other two sources. SO_2 is detected in G211 and G208. Among the four sources, G211 is the most line-rich target in which $\text{C}_2\text{H}_5\text{OH}$, NH_2CHO , and isotopologue species including $^{13}\text{CH}_3\text{OH}$, CH_2DOH , $^{13}\text{CH}_3\text{CN}$, and HC^{13}CCN were identified. G210 and G192 seem to have similar molecular composition. We consolidate and present in Table C1 the molecular line parameters, such as their rest frequencies, quantum numbers, excitation energies, etc., for the identified species and transitions.

Spectra of the molecule candidates were then modeled with the eXtended CASA Line Analysis Software Suite (XCLASS,⁶ Möller et al. 2017), which is a CASA toolbox for the molecular line synthesis. In myXCLASS, a function of XCLASS, each (gaseous) component of a molecule is assumed to be in its individual local-thermodynamic equilibrium (LTE). It means that the source function is in the form of Planck function with an rotational temperature (T_{rot}) which is expected to be the same as its kinetic temperature if all the levels are thermalized (Goldsmith & Langer 1999; Möller et al. 2017). To execute myXCLASS, users need to define the rotational temperature (T_{rot}), the size (θ_C), the col-

umn density (N), the local-standard-of-rest velocity (v_{LSR}) and the velocity width (Δv) of each component. These parameters are assumed to be identical for all transitions of each molecular component. By solving the radiative transfer equation for an isothermal object in one dimension with the given molecular component parameters, myXCLASS generates the synthetic spectrum and provides the transition list, intensities profiles and optical depth profiles of each component. We further employed Modeling and Analysis Generic Interface for eXternal numerical codes (MAGIX; Möller, T. et al. 2013), another CASA package, to optimizes these molecular component parameters within the given ranges through the XCLASS interface. XCLASS also provides the transition information based on the CDMS and JPL databases.

In our study, we fixed the component sizes (θ_C) to be the same as the deconvolved dust continuum (i.e. $\theta_C = \theta_S$) except for carbon monoxide C^{18}O , and optimized the other parameters using MAGIX. We note that this size may be overestimated for molecules (e.g. COMs) with compact emission and lead to underestimation of their column densities. For the extended molecules (e.g. carbon-chain molecules), this size may be underestimated but the column density averaged within the beam should be correctly estimated. Some molecules (e.g. hydrogen sulfide H_2S and sulfur monox-

⁶ <https://xclass.astro.uni-koeln.de/>

ide ^{34}SO), which are commonly detected in star-forming regions, were detected with only one line so that their (rotation) temperature could not be estimated and were fixed to be 100 K for their molecular column density calculation. We executed MAGIX for each molecule individually to find their correspondingly best-fit. With those numbers as inputs, we then executed a final run including all species for global optimization. Table 4 show the results from MAGIX optimization for G211, G208, G210 and G192, respectively. The parameter uncertainties estimated by the error estimation function of MAGIX are also presented. We also mark in Table 4 the molecules which are unresolved in their moment 0 map. The hydrocarbons and the carbon-chain molecules are resolved while saturated COMs are not, indicating that they are tracing the cold and the warm/hot regions, respectively. The S-bearing molecules are in general unresolved as well and they might be tracing the central region of the outflows due to their large linewidths. See Sect. 4 for more discussions.

The population diagram (or rotational diagram) is a common tool for estimating the rotational temperature, and the column density (Goldsmith & Langer 1999). Figure D1 shows the population diagram of CH_3OH in each source and the derived column densities and rotational temperatures are all consistent with the estimation of MAGIX. See Appendix D for the detail.

4. DISCUSSION

4.1. Source Overview

All four sources, except for G192, have been identified as YSOs in previous studies. G211, G208, and G210 were first identified as YSOs (MGM2012 518, 2433, and 777, respectively) in a mid-IR (spanning 3–24 μm) survey toward the Orion A and B clouds conducted with the Spitzer Space Telescope (Megeath et al. 2012). G211 and G210 were also identified as YSOs (FKV2013 771 and 593, respectively) in a survey from 0.4 to 24 μm toward the LYNDs 1641 cloud which is a star forming region in Orion A (Fang et al. 2013). G208, on the other hand, also coincides with OMC3/MMS6 (Chini et al. 1997). These three sources were further classified as Class 0 protostars (Furlan et al. 2016).

The three sources are also subjects of outflow studies in the literature (e.g. Manoj et al. 2016; van Kempen et al. 2016; Watson et al. 2016; Fischer et al. 2010; Velusamy et al. 2014). From our observations, the carbon monoxide CO $J=2-1$ images show a clear bi-conical outflow morphology in G211 and G210. (See Figure 1). In addition, silicon monoxide SiO is known as a good probe of shocks. The spatial distribution of the SiO $J=5-4$ emission in G211 and G210 is consistent with their outflows (See Figure 1.). Although no obvious outflow signature appears in our G208 CO $J=2-1$ image, Takahashi & Ho (2012) and Takahashi et al. (2019) have detected a very compact ($< 5''$) bipolar molecular out-

flow in the CO $J=3-2$ and $2-1$ lines, respectively, indicating the object being in an extremely young stage of formation. Our detection of compact SiO emission is consistent with what was inferred by Takahashi & Ho (2012).

G192 is the only object in λ Orionis among the four sources. While it has not been associated with any known YSOs in the literature, our observations indicate the presence of bi-conical outflow in its CO $J=2-1$ emission.

4.2. Hot Corino Identification

A hot corino is normally identified according to their compact (< 100 au), abundant and presumably passively thermal-evaporated COMs (Herbst & van Dishoeck 2009). Methanol (CH_3OH), as the root of the COMs in the grain origin scheme in particular, is considered as an important indicator of the presence of hot corinos. Toward all the four sources (G211, G210, G192, and G208), we detected CH_3OH with very compact distribution (See Figure 1). Similar to the continuum analysis, we tried to use CASA 2D Gaussian fitting to fit the CH_3OH integrated intensity maps. The distributions of CH_3OH emission, even for the transition with the low excitation energy ($E_u \sim 55$ K) we have, are very compact as the fitting routine suggested the sources to be unresolved. Additionally, the high rotational temperatures (> 165 K) are also consistent with and indicative of thermal evaporation of icy CH_3OH . Unlike those, often extended, CH_3OH emission seen in cold, dark clouds (Friberg et al. 1988), cold, massive prestellar clumps (Sanhueza et al. 2013), or outflow (Araya et al. 2008), the evidences we present suggest the four objects be hot corinos. Further studies with higher angular and spectral resolutions would help to solidify (or falsify) this classification.

Both COMs and carbon-chain molecules are detected in our four sources. Lefloch et al. (2018) proposed a classification scheme based on the ratio of the number of O-bearing molecular species over that of the hydrocarbon species. While the scheme was not related to any chemical models but rather phenomenological, the definitions of the O-bearing and hydrocarbon molecules are in the type of $\text{C}_x\text{H}_y\text{O}_z$ and C_xH_y , respectively. In our samples, the values of this ratio are all larger than 1.9, which is the lower limit suggested by Lefloch et al. (2018) for a hot corino. Note that this classification is the N- and S- bearing molecules are not involved in the definitions.

Table 5 shows the fractional column density, X , of COMs relative to hydrogen. The fractional column densities of methanol $X(\text{CH}_3\text{OH})$ are larger than 10^{-8} , which is typical of a hot corino, except in G208. In addition to methanol, there are more COMs (e.g. acetaldehyde CH_3CHO and methyl formate HCOOCH_3) detected in G211 and G208. In G208, $X(\text{CH}_3\text{CHO})$ and $X(\text{HCOOCH}_3)$ are 4.1×10^{-10} and 6.6×10^{-10} , respectively, and they are 5 and 20 times larger,

Table 4. Molecule component list calculated by MAGIX.

Formula	G211.47-19.27S				G208.68-19.20N1				G210.49-19.79W				G192.12-11.10			
	T_{rot} K	N cm^{-2}	v_{LSR} km s^{-1}	Δv km s^{-1}	T_{rot} K	N cm^{-2}	v_{LSR} km s^{-1}	Δv km s^{-1}	T_{rot} K	N cm^{-2}	v_{LSR} km s^{-1}	Δv km s^{-1}	T_{rot} K	N cm^{-2}	v_{LSR} km s^{-1}	Δv km s^{-1}
CH ₃ OH	185±4	8.5E+16 (0.1) [†]	2.2	6.2	298±3	4.9E+15 (0.1) [†]	10.4	4.1	155±3	7.1E+15 (0.1) [†]	8.1	9.8	173±5	7.2E+15 (0.2) [†]	9.6	9.2
¹³ CH ₃ OH	156±6	9.2E+15 (0.1) [†]	2.0	7.7
CH ₂ DOH	67±5	2.3E+16 (0.1) [†]	1.8	6.0
CH ₃ CHO	127±4	1.4E+15 (0.1) [†]	2.4	4.9	223±5	5.6E+14 (0.1) [†]	11.0	2.8
CH ₃ OCHO	197±6	8.5E+15 (0.1) [†]	3.3	7.8	143±3	1.2E+15 (0.1) [†]	11.4	1.7
C ₂ H ₅ OH	304±6	1.7E+16 (0.1) [†]	1.5	7.2
¹³ CH ₃ CN	295±6	8.2E+13 (0.1) [†]	2.4	7.0
CCD	62±6	2.3E+14 (0.1)	4.8	2.2	24±3	8.1E+13 (0.1)	11.1	1.4	24±3	4.3E+13 (0.1)	8.7	2.0	19±4	5.5E+13 (0.2)	10.3	2.3
c-C ₃ H ₂	32±4	1.2E+14 (0.1)	5.1	3.1	22±3	5.7E+13 (0.1)	11.0	1.0	20±3	2.2E+13 (0.1)	8.6	1.7	13±5	3.7E+13 (0.2)	10.4	1.8
H ₂ CO	189±4	8.2E+15 (0.1)	4.6	4.9	39±3	5.9E+14 (0.1)	11.1	1.6	51±3	2.7E+14 (0.1)	8.4	2.6	45±5	2.9E+14 (0.2)	10.5	2.8
D ₂ CO	136±5	3.5E+14 (0.1) [†]	3.9	4.6	38±3	5.7E+13 (0.1)	11.2	1.9	40±3	2.4E+13 (0.1)	8.4	2.2	22±5	3.3E+13 (0.2)	10.5	2.3
DCO ⁺	100	2.6E+13 (0.2)	4.8	2.4	100	2.9E+13 (0.1)	11.3	1.8	100	9.1E+12 (0.1)	8.3	1.3	100	1.8E+13 (0.2)	10.4	2.0
N ₂ D ⁺	100	1.9E+12 (0.1)	11.9	1.4
HCOOH	100	9.2E+14 (0.1) [†]	1.5	4.3	100	3.5E+13 (0.1) [†]	10.8	0.8	100	7.7E+13 (0.1) [†]	8.1	6.5	100	1.2E+14 (0.2) [†]	9.9	7.2
DCN	100	4.0E+13 (0.1)	4.3	4.8	100	1.2E+13 (0.1)	11.1	1.7	100	7.3E+12 (0.1)	8.5	3.1	100	6.8E+12 (0.2)	10.5	3.2
HCCCN	100	5.0E+13 (0.1)	4.3	5.3	100	1.3E+13 (0.1) [†]	11.3	3.4	100	7.1E+12 (0.1) [†]	8.2	4.9	100	6.0E+12 (0.2) [†]	11.0	5.3
HC ¹³ CCN	100	1.6E+13 (0.2) [†]	5.0	5.9
HNCO	205±4	1.2E+15 (0.1) [†]	2.5	7.8	161±4	1.2E+14 (0.1) [†]	8.4	10.0	100	3.8E+13 (0.2) [†]	10.8	5.5
NH ₂ CHO	295±6	3.6E+14 (0.1) [†]	2.5	7.7
H ₂ S	100	2.1E+15 (0.1) [†]	3.2	5.8	100	1.6E+15 (0.1) [†]	11.2	1.9	100	1.7E+14 (0.1) [†]	9.2	3.8	100	3.5E+14 (0.2) [†]	10.6	6.7
¹³ CS	100	5.5E+13 (0.1) [†]	4.3	5.3	100	3.2E+13 (0.1)	11.1	1.8	100	3.6E+12 (0.1) [†]	8.4	2.1	100	6.3E+12 (0.2) [†]	10.4	3.4
OCS	70±4	2.1E+15 (0.1) [†]	3.0	5.8	122±3	4.3E+14 (0.1) [†]	11.5	2.1	127±6	2.3E+14 (0.1) [†]	8.5	7.3	86±5	2.4E+14 (0.2) [†]	10.5	7.7
³⁴ SO	100	2.1E+14 (0.1) [†]	4.8	6.9	100	6.8E+13 (0.1)	11.2	2.8	100	3.8E+13 (0.1) [†]	11.5	13.9	100	2.8E+13 (0.2) [†]	10.6	3.1
SO ₂	100	3.0E+15 (0.2) [†]	4.1	12.7	100	6.2E+14 (0.1) [†]	10.2	9.6
SiO	100	5.9E+13 (0.1) [†]	8.9	13.1	100	4.3E+13 (0.1) [†]	30.2	40.5	100	3.3E+12 (0.1) [†]	8.9	5.1
C ¹⁸ O	100	1.3E+16 (0.1)	5.0	3.2	100	2.9E+13 (0.1) [†]	-16.1	16.7	100	1.4E+16 (0.1)	8.3	1.7	100	1.9E+16 (0.2)	10.5	2.4

NOTE— The T_{rot} is the rotational temperature and the $T_{\text{rot}}=100$ K without any error interval is fixed in the MAGIX simulation. The N is the column density and the value in the brackets is their standard deviations in $\log_{10} N$. Their values 0.1-0.2 correspond to fractional errors ranging around 25%-70%. The v_{LSR} is the local standard of rest velocity and the Δv is the velocity width. Their standard deviation values are around 0.4 km s^{-1} . The component size of C¹⁸O is assumed to be $8''$. The SiO profile in G208 was assumed to be constituted of two components. The daggers mark the unresolved molecules (i.e. observed source sizes is smaller than the synthesized beam size on the moment 0 map) and note that the column densities of these spatially unresolved molecules may be underestimated.

Table 5. The fractional column density with respect to H₂ of the COMs in hot corinos.

$\times 10^{-9}$	CH ₃ OH	CH ₃ CHO	HCOOCH ₃	HNCO	HCOOH	NH ₂ CHO	C ₂ H ₅ OH	Ref.
G211.47-19.27S	110	2.09	13.1	1.64	1.32	0.51	26.9	
G208.68-19.20N1	2.92	0.41	0.66	...	0.03	
G210.49-19.79W	64.9	1.16	0.76	
G192.12-11.10	43.5	0.26	0.83	
IRAS 16293-2422 B	833	10.0	21.7	...	4.67	...	19.2	1, 2
B335	380	2.40	4.60	17.0	4.70	0.40	3.80	3
HH-212	160	3.90	8.40	...	5.3	0.42	7.10	4

NOTE— This fractional column density of CH₃OH may be affected by its opacity (See Sect. 4.2).

References— 1. Jørgensen et al. (2018); 2. Jørgensen et al. (2016) 3. Imai et al. (2016); 4. Lee et al. (2019)

Table 6. The fractional column density of COMs relative to CH₃OH, [COM]/[CH₃OH].

$\times 10^{-3}$	CH ₃ CHO	HCOOCH ₃	C ₂ H ₅ OH	Ref.
G211.47-19.27S	18.9	119	243	
G211.47-19.27S [†]	2.0	12.9	26.4	
G208.68-19.20N1	141	225	...	
IRAS 16293-2422 B	12.0	26.0	23.0	1
B335	6.32	12.1	10.0	2
HH-212	24.4	52.5	44.4	3

NOTE—[†] The column density of CH₃OH is derived from the column density of ¹³CH₃OH assuming ¹²C/¹³C ratio is 70.

References—1. Jørgensen et al. (2018); 2. Imai et al. (2016); 3. Lee et al. (2019)

respectively, in G211. Ethanol C₂H₅OH is also detected in G211 and $X(\text{C}_2\text{H}_5\text{OH})$ is 2.69×10^{-8} , comparable to the value in IRAS 16293-2422 B (Jørgensen et al. 2018, see Table 5). The brightness temperatures of all these COM emissions are well below < 1 K, implying these emissions are optically thin. However, if some of these lines suffer severely from beam dilution, their opacities could be higher, which will lead to higher fractional column densities.

Table 6 shows the fractional column densities of COMs with respect to CH₃OH ([COM]/[CH₃OH]). In IRAS 16293-2422 B, B335, and HH-212, [CH₃CHO]/[CH₃OH] is roughly half of [HCOOCH₃]/[CH₃OH] and [C₂H₅OH]/[CH₃OH]. In our sources, CH₃CHO and HCOOCH₃ are detected in G211 and G208, and C₂H₅OH is detected in G211. The ratio [CH₃CHO]/[CH₃OH] in G211 is comparable to B335 and HH-212; however, [CH₃CHO]/[CH₃OH], [HCOOCH₃]/[CH₃OH], and [C₂H₅OH]/[CH₃OH] in G211 and G208 (without C₂H₅OH) are an order of magnitude larger than the typical values in other hot corinos.

The discrepancy in these ratios could result from a few possibilities. First, it may be due to different beam dilution factors of different COM emissions. In the current analyses, we have assumed that all COM emissions originate from the same extent as the dust continuum. As noted earlier, the COM emissions in fact appear to be more compact than the continuum emission. The COM emission size may have been over-estimated, leading to underestimations of the COM column densities. We speculate a more compact CH₃OH emission size (than the continuum as well as HCOOCH₃ and C₂H₅OH), which may correspond to a higher CH₃OH column density and hence reduces [HCOOCH₃]/[CH₃OH], and [C₂H₅OH]/[CH₃OH]. Our current spatial resolution is not yet sufficient to fully resolve the region from where the COM

emission is emitted. Future investigations using the 12m ALMA data may help to validate this scenario.

Second, CH₃OH abundances in our sources may have been underestimated due to opacity. A hint of this is the column density ratio of methanol and its isotopologues. In G211, the ratio of methanol isotopologues ¹²C/¹³C derived from [CH₃OH]/[¹³CH₃OH] is ~ 10 , which is lower than the typical ¹²C/¹³C ~ 70 in the local interstellar medium (ISM) (Wirstrom et al. 2011). The ratio appears comparable to [CH₃OH]/[¹³CH₃OH] of other Class 0 hot corinos, e.g. 24 in HH-212 (Lee et al. 2019, see Table 7). CH₃OH transitions with high optical depths have been observed in Taquet et al. (2015) and Lee et al. (2019) for the cases of IRAS 2A and HH-212, respectively. If we assume that CH₃OH column density in G211 is seven times larger in order to make ¹²C/¹³C compatible with the local ISM, the high [COM]/[CH₃OH] observed in our sources will be reduced. In this case, however, the deuteration ratios of CH₃OH ([CH₂DOH]/[CH₃OH], Table 7) and [CH₃CHO]/[CH₃OH] become small as compared to others. Instead, we can estimate those [COM]/[CH₃OH] ratios in G211 by utilizing the column density of ¹³CH₃OH as a proxy of CH₃OH, assuming ¹³CH₃OH emission is optically thin and a ¹²C/¹³C of 70. As shown in Table 6, the [COM]/[CH₃OH] ratios in G211 appear comparable to the other hot corino sources.

Finally, Bergner et al. (2019) compared the [COM]/[CH₃OH] ratio in hot corinos Ser-emb 1, Ser-emb 8, and Ser-emb 17 with other hot corinos and found that it spanned two orders of magnitude. They suggested this variation is the result of the local environment and/or the time-dependent warm-up chemistry.

The only COM detected in G210 and G192 is CH₃OH. Assuming other COMs such as CH₃CHO and HCOOCH₃ bear similar [COM]/[CH₃OH], their column densities would be below $\sim 10^{14}$ cm⁻² and become hardly detected as is the case we have.

4.3. Formaldehyde (H₂CO)

Once CO is frozen onto grain surfaces, formaldehyde (H₂CO) can be formed via hydrogen addition reactions from CO and then further forms CH₃OH, and their photodissociated radicals are parents of some COMs synthesized via both ice and gas-phase chemistry (Herbst & van Dishoeck 2009). H₂CO and CH₃OH are detected in all our four sources and they seem to show similarly compact distributions (See Fig. 1). Although H₂CO and CH₃OH are detected in similar regions, the excitation temperatures of the two species are different in G208, G210 and G192. Assuming that their kinetic temperatures are the same, it indicates that their energy levels may not be thermalized.

H₂CO and its isotopologue, double-deuterated formaldehyde (D₂CO), are both detected in all our hot corino

sources. Their rotational temperatures are intermediate among the species detected with more than one transition and hence with temperatures estimated from XCLASS, and their linewidths are intermediate as well.

4.4. Formamide (NH_2CHO)

Studies show that formamide NH_2CHO is a potential key species in pre-biotic evolution. It has been proposed as one of the main components of both (pre)genetic and (pre)metabolic processes (Saladino et al. 2012). In G211, the fractional abundance with respect to H_2 of formamide ($X(\text{NH}_2\text{CHO})$) and of isocyanic acid ($X(\text{HNCO})$) are 1.3×10^{-10} and 4.1×10^{-10} , respectively. They are within the ranges (from 10^{-11} to 10^{-9} and from 10^{-12} to 10^{-8} , respectively) in star-forming regions with H_2CO detections (Lpez-Sepulcre et al. 2015).

The $[\text{NH}_2\text{CHO}]/[\text{HNCO}]$ in G211 is about 0.32, which is comparable to 0.14 in the hot corino L483 (Oya et al. 2017). Furthermore, it was suggested that NH_2CHO is chemically related to HNCO and their abundances follow a tight correlation, $X(\text{NH}_2\text{CHO}) = 0.04 X(\text{HNCO})^{0.93}$ in star-forming regions with H_2CO detection (Lpez-Sepulcre et al. 2015). $X(\text{NH}_2\text{CHO})$ in G211 is 5.1×10^{-10} , which is comparable to the value estimated by the above formula (2.7×10^{-10}). The formation mechanism of NH_2CHO supported by this tight correlation is the hydrogenation of HNCO in icy grain mantles (Raunier et al. 2004; Garrod et al. 2008; Jones et al. 2011; Lpez-Sepulcre et al. 2015). However, a recent laboratory experiment suggested that this pathway of NH_2CHO forming from the hydrogenation of HNCO is insufficient (Noble et al. 2015).

Another mechanism of NH_2CHO formation is via the reactions between NH_4^+ and H_2CO in the gas-phase: (1) the radiative association reaction and the followed dissociative recombination of NH_2CHO^+ (Quan & Herbst 2007), (2) the ionmolecule reaction and the followed subsequent electron recombination of NH_3CHO^+ (Halfen et al. 2011) and (3) the reaction between amidogen and formaldehyde ($\text{NH}_2+\text{H}_2\text{CO}$, Barone et al. 2015). In HH-212, the NH_2CHO and the D_2CO have similar spatial distributions. It indicates the possible correlation between NH_2CHO and H_2CO and further supports the formation of NH_2CHO from H_2CO in gas-phase (Lee et al. 2017). Due to the limitation of the spatial resolution, we are not able to examine the spatial correlation between them.

Finally, a laboratory experiment show that NH_2CHO is also possibly formed via the icy mixtures of carbon monoxide and ammonia (i.e. $\text{CO}:\text{NH}_3$) irradiated by energetic particles such as electrons (Jones et al. 2011). Moreover, the UV photoprocessing of $\text{CO}:\text{NH}_3$ and $\text{CO}:\text{CH}_4$ ice samples produces NH_2CHO and CH_3CHO , respectively, and the predicted abundance of NH_2CHO is 2 to 16 times larger than

CH_3CHO (Martin-Domenech et al. 2020). It is opposite to the current observations via the other hot corinos as well as our studies (e.g. ~ 0.1 in HH-212, Lee et al. 2019) so an additional chemical mechanism may be required.

In G210 and G192, $X(\text{HNCO})$ are 1.2×10^{-9} and 2.6×10^{-10} , respectively, which are within the range presented by Lpez-Sepulcre et al. (2015). Their $X(\text{NH}_2\text{CHO})$ are estimated to be $\sim 10^{-10}$ and their column densities of NH_2CHO are therefore estimated to be in the order of $\sim 10^{13}$, which may be insufficient to be detected.

4.5. Hydrocarbons and Long Carbon-Chain Molecules

Deuterated ethynyl (CCD), cyclopropenylidene ($\text{c-C}_3\text{H}_2$) and cyanoacetylene (HCCCN and HC^{13}CCN in G211), which bear carbon-chains or carbon-rings, are also detected. The estimated rotational temperatures of CCD and $\text{c-C}_3\text{H}_2$ are distinctly cooler: ~ 25 K and ~ 15 K, respectively, in G208, G210 and G192, and slightly warmer (~ 63 K and ~ 32 K, respectively) in G211. It is consistent with the trend that temperatures are warmer in G211 for all the species. The cooler temperatures of these species are also compatible with their more extended emission presented in Fig. 1. Meanwhile, their linewidths are narrow ($\Delta v < 2.5$ km s^{-1} in 208, G210 and G192 and $\Delta v < 3.5$ km s^{-1} in G211) compared with the linewidths of the COMs and the outflow tracers.

HCCCN, a long carbon-chain molecule, is detected in all four sources, but its linewidths are wider than that of CCD and $\text{c-C}_3\text{H}_2$. Assuming the rotational temperature of HC^{13}CCN to be 100K, the column density of HCCCN is in general higher than that of $\text{c-C}_3\text{H}_2$ and CCD within an order of magnitude.

The HC^{13}CCN , which is an isotopologue of HCCCN, is detected in G211 only. Under the assumption of $T_{\text{rot}} = 100$ K, the ratio $[\text{HC}^{13}\text{CCN}]/[\text{HCCCN}] \sim 0.33$. While here we assumed an excitation temperature higher than the calculated value of the other carbon-chain molecule (i.e. CCD), the derived isotopic ratio is meaningful. This is because the column density or abundance ratios will be simply very close to the intensity ratio of the $J=24-23$ (with similar molecular parameters) seen in both isotopologues. It is an order of magnitude higher than what was reported by Araki et al. (2016) in the low-mass star-forming region L1527.

4.6. Deuterations

The deuterium fractionation (D/H ratio) has been suggested as an indicator of the gas temperature in (low mass) star formation. This D/H ratio is in general anti-correlated with the temperature during the cold prestellar phase (Roberts & Millar 2000; Persson et al. 2018). The deuterium fractionation then decreases through the protostellar evolution (Taquet et al. 2014; Bianchi et al. 2017). Our measurements of D/H from H_2CO appear to follow this trend

Table 7. The column densities ratio between the isotopes of the formaldehyde and methanol in other hot corinos.

	$\frac{[D_2CO]}{[H_2CO]}$	$\frac{[CH_2DOH]}{[CH_3OH]}$	$\frac{[CH_3OH]}{[^{13}CH_3OH]}$	Ref.
G211.47-19.27S	0.043	0.27	9.2	
IRAS 16293-2422 B	0.006	0.37	...	1
HH-212	...	0.12	24	2
IRAS 4B	0.046	0.43	...	3
IRAS 2A	0.052	0.62	26	3, 4

References—1. López-Sepulcre et al. (2017); 2. Lee et al. (2019); 3. Parise et al. (2006); 4. Taquet et al. (2015);

since the $[D_2CO]/[H_2CO]$ of the hottest source (G211) is the smallest. The $[D_2CO]/[H_2CO]=0.039, 0.091, 0.089$ and 0.112 for G211, G208, G210 and G192, respectively. These values are comparable to those of other hot corinos (e.g. 0.046 in IRAS 4B; Parise et al. 2006) except IRAS 16293-2422 B, which is an order of magnitude lower than others (see Table 7). Since the temperature of H_2CO in IRAS 16293-2422 B is higher (~ 107 K; Persson et al. 2018) compared to the sources we presented, it seemingly suggests that the evolution stage of the sources we presented is in general earlier than that of IRAS 16293-2422 B, which is estimated to be $\sim 10^5$ yr (Persson et al. 2018). We note, however, that deuterated methanol is also detected in G211, and the D/H ratio of methanol ($[CH_2DOH]/[CH_3OH]$) is 0.27, slightly lower than that of IRAS 16293-2422 B. This trend is the opposite to the case for formaldehyde. The use of D/H for diagnosing or differentiating evolutionary ages of objects thus needs to be applied with caution. We note that correcting for the optical thickness of CH_3OH may further reduce the derived $[CH_2DOH]/[CH_3OH]$ values.

4.7. S-bearing Molecules

Several S-bearing molecules are commonly detected toward the four sources. They mostly have wide linewidths except for carbon monosulfide ^{13}CS . The abundance ratio $[^{13}CS]/[H_2S]$ is ~ 0.02 in all the four sources. In Drozdovskaya et al. (2018) and Le Gal et al. (2019), the abundance ratio of sulfur-bearing molecules with respect to CS in IRAS 16293-2422 B and other environments was presented but the uncertainties made it difficult to make the comparison.

The formation or synthesis processes of those S-bearing molecules remains unclear. Models proposed that sulfur monoxide SO and sulfur dioxide SO_2 are formed in the gas phase from hydrogen sulfide H_2S and carbonyl sulfide OCS evaporated from grain mantles (Palumbo et al. 1997; Hatchell et al. 1998; Esplagues et al. 2014).

Recently, Luo et al. (2019) presented opposite temperature dependencies in the abundances of carbon-sulfur compounds (e.g. ^{13}CS and OCS) and carbon-free sulfur-bearing species (e.g. H_2S , ^{34}SO and SO_2) in the Orion KL region. Unfortunately our sample of four hot corinos is too small to investigate these effects within the Orion Molecular Cloud and larger observational samples are required (Le Gal et al. 2019).

4.8. Trends and Comparisons

Toward the four targets, tracers like CO and SiO show the broadest linewidths, clearly marking energetic outflow activities. While COMs emission are not spatially resolved, their linewidths are also wide and it is especially true for CH_3OH . Lee et al. (2017) and Lee et al. (2019) imaged with ALMA a set of COM emission from the atmosphere of a Keplerian rotating disk around the central YSO. All four sources we are studying, as alluded earlier, are associated with molecular outflows. They therefore could have circumstellar disks that mediate the accretion and the launching of the outflows. It is tempting to speculate if the COMs we detected are of a similar origin. If the linewidths are really as a result of rotating motion in the putative circumstellar disks, the broader linewidths of CH_3OH suggest that it could originate from the inner part of an incipient disk, as compared to, for example, H_2CO . This is in general consistent with the rotational temperature of CH_3OH being higher than that of H_2CO . On the other hand, the very presence of the outflows also implies the alternative possibility that our observed COMs like CH_3OH are in fact related to the shock activities at the very base of the molecular outflows in the vicinity of the central YSO.

The hydrocarbon and long carbon-chain species such as CCD, $c-C_3H_2$, and HCCCN have relatively lower rotational temperatures and narrower linewidths. Such characteristics, which are similar to those of CCH in the hot corino B335 (Imai et al. 2016), indicate that these species are more extended as compared to the continuum and COMs (e.g. $c-C_3H_2$ in figure 1). The line profile of CCH in B335 in fact hinted at the existence of double-peak signatures due to absorption by the cold envelope gas. We unfortunately are not able to discriminate such features in our study at this stage due to the limited spatial and spectral resolution of the data analyzed. Another species that in general bears a narrow linewidth is formylium DCO^+ .

Among the four hot corino objects, the line widths (Δv) of molecular emission in G208 are overall narrower than those of the other three sources. Taking CH_3OH as an example, Δv is ~ 7.5 $km\ s^{-1}$ in G211, G210 and G192. In contrast, Δv in G208 is 2.3 $km\ s^{-1}$. Meanwhile, the lower bolometric temperature in G208 would indicate a smaller region where CH_3OH can get evaporated from the icy mantle. The lower velocity at a closer distance may imply a lower mass for the

central protostellar object in G208, if CH_3OH emission is tracing rotational motion in the circumstellar disk. Indeed, the relatively low bolometric luminosity, the most compact and likely young molecular outflows, and the highest envelope mass are all qualitatively but coherently suggestive of a younger stage. Alternatively, the compact morphology of the CO outflow emission and the narrower linewidths of trace molecules like CH_3OH may be resulted from an outflow-disk/envelope system viewed pole-on. Finally, there is also the possibility that its COM emissions have a different origin instead of tracing a genuine hot corino like the other three objects, hence giving a different $[\text{COM}]/[\text{CH}_3\text{OH}]$.

In our ALMA observation, 72 fields were observed with 48 sources being of protostellar nature for their associations with YSOs and/or molecular outflows (Dutta et al., in preparation). Among those, we have identified 4 hot corino sources, a fraction of $\sim 8\%$ and the three hot corinos found in Orion A are Class 0 HOPS objects. We note that the 8% face value needs to be treated with caution as the sample size is limited and the selection of the initial 72 fields is based on the JCMT SCUBA-2 detections in the PGCC targets, which implies that these sources are likely more embedded in large scale envelopes. Other investigations of similar kind show different occurrence rates of corino signatures. For example, the observations with ALMA, as part of the recent VLA/ALMA Nascent Disk and Multiplicity (VAN-DAM) survey, targeted toward four fields in the Orion Molecular Cloud-2 (OMC2) Far-infrared sources OMC2-FIR4 and OMC2-FIR3, identified eleven $870 \mu\text{m}$ continuum sources. Among them, eight sources are associated with HOPS objects and two of them, HOPS 108 in OMC2-FIR4 and HOPS 370 in OMC2-FIR3, are likely harboring hot corinos as well (Tobin et al. 2019). The Continuum And Lines in Young ProtoStellar Objects (CALYPSO) program by IRAM surveyed 16 Class 0 protostellar systems, some with multiple protostars, in nearby clouds and detected compact (nearly all < 100 au) CH_3OH emission around about half of the protostellar objects (Belloche et al. 2020). It further suggested a luminosity threshold of $4 L_\odot$, above which YSOs in their sample exhibit spectral features from at least one COM, and no COM emission was detected for sources with luminosities lower than $2 L_\odot$. They attributed this, though, to the sensitivity of the observations instead of the intrinsic property of the sources with lower luminosities and noted that there exist other COMs-detected sources fainter than $2 L_\odot$ (e.g. $\sim 0.7 L_\odot$ in B1-bS and $\sim 0.7 L_\odot$ in B335 Hirano & Liu 2014; Evans et al. 2015). In our study, three hot corinos, of which the SEDs have been established by Furlan et al. (2016), have luminosities well above that threshold. The luminosity of the fourth source (G192) is at a level of $\sim 14 L_\odot$ (Dutta et al., in preparation). There are, on the other hand, protostellar objects in our sample with luminosities much greater than 4

L_\odot but show no clear sign of hot corinos. It remains unclear whether hot corino is an evolutionary stage that YSOs would generally experience, or only YSOs with certain properties or in certain environments would develop into hot corinos. A full census of YSOs in different environments and evolutionary stages will help in revealing the prevalence of hot corinos and addressing the above questions on a statistical basis.

5. CONCLUSIONS

1. Based on a survey toward selected Planck Galactic Cold Clump (PGCC) samples in Orion Molecular Cloud Complex nebula A and B and λ Orionis, we report the detection of four hot corinos based on the presence of warm, compact CH_3OH emission in the four sources, and additional COMs in two of them. Given the positional association of their 1.3 mm continuum, three of these four sources are identified as Class 0 young stellar objects associated with HOPS sources, and the remaining one is for the first time identified as a protostellar core. Further studies with higher angular and spectral resolutions would help to solidify (or falsify) the classification of hot corinos.
2. Some fractional column densities of COMs with respect to methanol, $[\text{COM}]/[\text{CH}_3\text{OH}]$, are an order of magnitude higher than those in other hot corinos in the literature. In G211.47-19.27S, the $^{12}\text{C}/^{13}\text{C}$ ratio of methanol ($[\text{CH}_3\text{OH}]/[^{13}\text{CH}_3\text{OH}]$) is ~ 10 which is comparable to the values in other hot corinos in the literature. This $^{12}\text{C}/^{13}\text{C}$ ratio is lower than is typical in the ISM (~ 70). The estimated column densities could be affected by beam dilution since the COM emission is not resolved. In addition, the relative abundances with respect to methanol, as well as the $^{12}\text{C}/^{13}\text{C}$ ratio in CH_3OH may be affected by its opacity. Variations between hot corinos are in any case expected due to the chemical diversity found in the literature (Bergner et al. 2019).
3. In G211.47-19.27S, both NH_2CHO , a pre-biotic molecule, and HNCO are detected. Their estimated column densities are $\sim 3.6 \times 10^{14} \text{ cm}^{-2}$ and $\sim 1.2 \times 10^{15} \text{ cm}^{-2}$, respectively. The abundance ratio $[\text{NH}_2\text{CHO}]/[\text{HNCO}]$ is about twice of the estimation based on Lpez-Sepulcre et al. (2015). HNCO is also detected in G210.49-19.79W and G192.12-11.10 without the detection of NH_2CHO . The lack of detection of NH_2CHO may result from the limitation of its low abundance.
4. The D/H ratios of formaldehyde ($[\text{D}_2\text{CO}]/[\text{H}_2\text{CO}]$) of these four sources are in general negatively correlated to their excitation temperatures. These

Figure 1. The moment 0 images of (from left to right, from top to bottom) continuum, CO (2-1, $f_{\text{rest}}=230538$ MHz, $E_u=16$ K), SiO (5-4, $f_{\text{rest}}=217105$ MHz, $E_u=31$ K), CH₃OH (5(1,4)-4(2,3) E, $f_{\text{rest}}=216946$ MHz, $E_u=55$ K), H₂CO (3(2,2)-2(2,1), $f_{\text{rest}}=218476$ MHz, $E_u=68$ K) and c-C₃H₂ (3(3,0)-2(2,1), $f_{\text{rest}}=216279$ MHz, $E_u=68$ K). The contours represent the continuum at [5,10,20,40,60,80] σ_{CONT} where σ_{CONT} is the RMS noise of each continuum (see Table 2). The color scale of the continuum and the other moment 0 images are in the unit of Jy beam⁻¹ and Jy beam⁻¹ km s⁻¹. The magenta marker is the location of each corresponding HOPS object (see Table 3). Note that there is no SiO $J=5-4$ transition detection in G192.12-11.10.

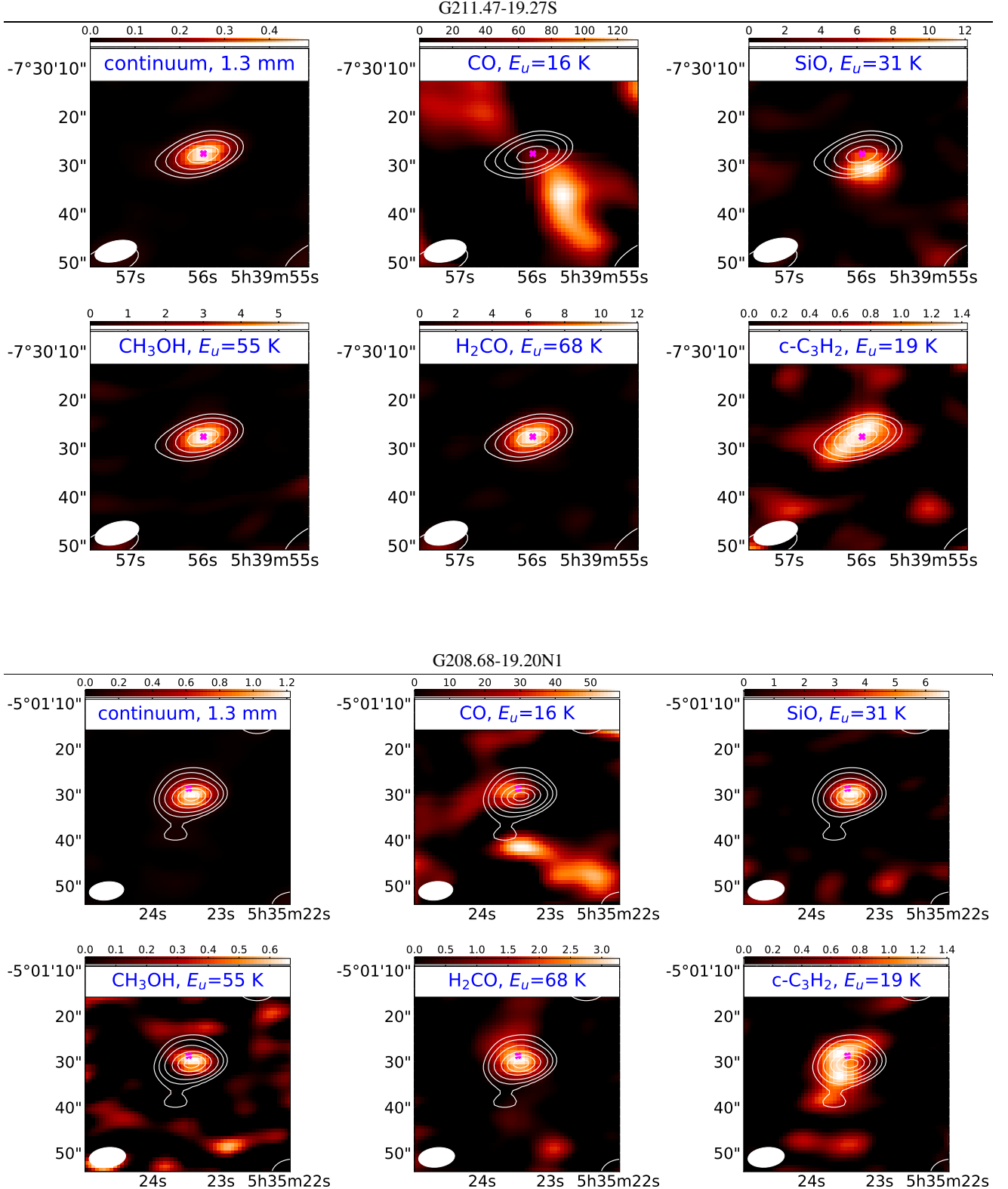
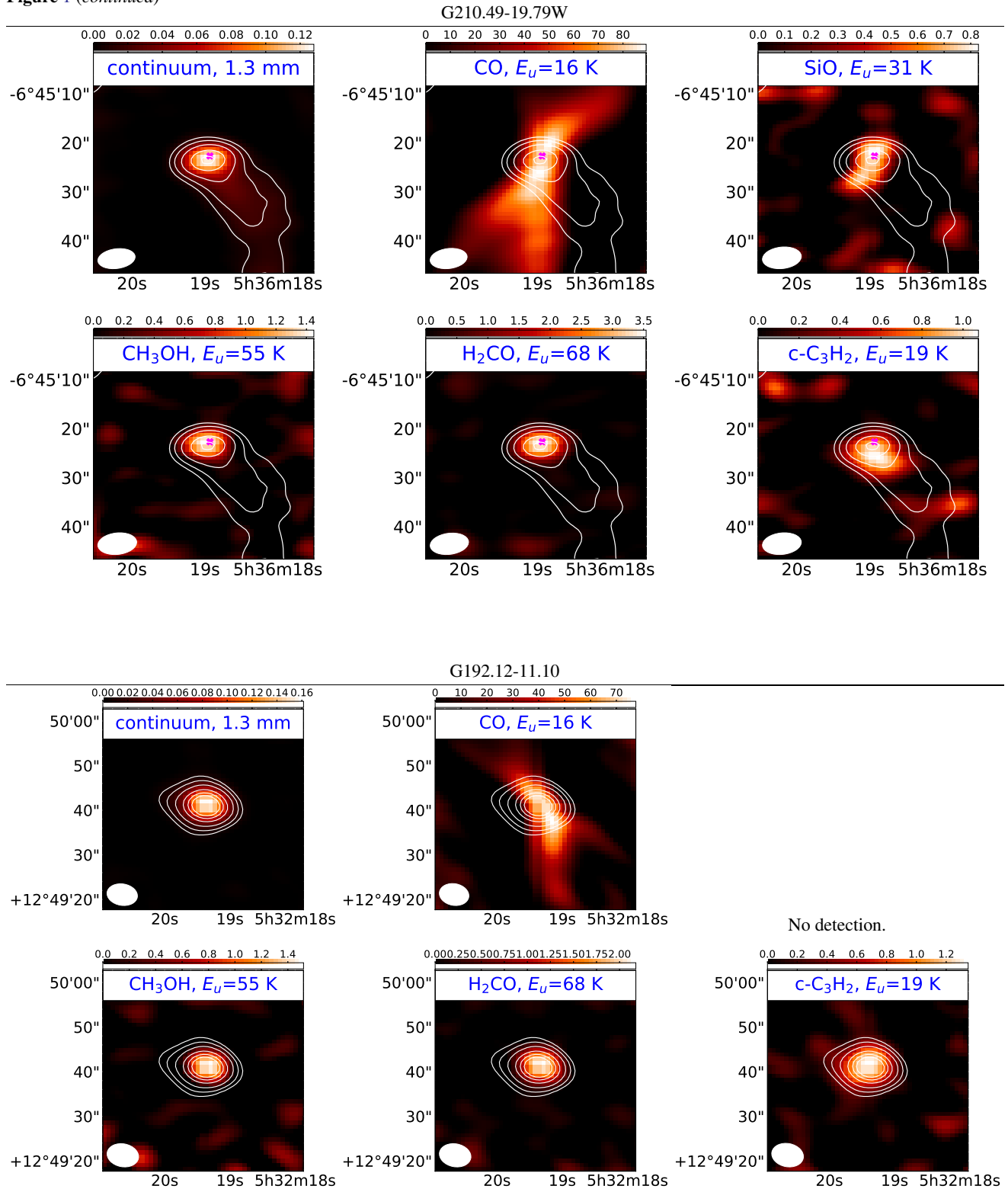


Figure 1 (continued)



$[D_2CO]/[H_2CO]$ as well as the the D/H ratio of methanol ($[CH_2DOH]/[CH_3OH]$) in G211.47-19.27S are comparable to those of other hot corinos in the literature except the $[D_2CO]/[H_2CO]$ in IRAS 16293-2422 B. This may result from the difference of either their temperatures or their evolutionary stages. Note that the D/H ratio of methanol ($[CH_2DOH]/[CH_3OH]$) is also possibly affected by the optical thickness of CH_3OH emission.

5. The hydrocarbons, $c-C_3H_2$ and CCD, are also detected in the four sources. They have cooler excitation temperatures and show more extended spatial distributions compared to COMs. The HCCCN is detected in all the four sources and its isotope, $HC^{13}CCN$ is detected in G211.47-19.27S. The ratio $[HC^{13}CCN]/[HCCCN]$ is an order of magnitude higher than the value in the low-mass star forming region L1527.

6. About 8% of the protostellar objects in our survey are identified as hot corinos. A more complete study of protostellar cores is required to reveal whether hot corino is a general stage of low-mass star formation.

ACKNOWLEDGMENTS

This paper makes use of the following ALMA data: ADS/JAO.ALMA#2018.1.00302.S. ALMA is a partnership of ESO (representing its member states), NSF (USA) and NINS (Japan), together with NRC (Canada), MOST and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO and NAOJ. SYH and SYL acknowledge support from the Ministry of Science and Technology (MoST) with grants 108-2112-M-001-048- and 108-2112-M-001-052-. N. Hirano acknowledges support from the Ministry of Science and Technology (MoST) with grant 108-2112-M-001-017. We thank Neal J. Evans and Siyi Feng for their useful comments in improving the manuscript. JHe thanks the National Natural Science Foundation of China under grant Nos. 11873086 and U1631237 and support by the Yunnan Province of China (No.2017HC018). This work is sponsored (in part) by the Chinese Academy of Sciences (CAS) through a grant to the CAS South America Center for Astronomy CASSACA) in Santiago, Chile. C.W.L. is supported by the Basic Science Research Program through the National Research Foundation of Korea (NRF) funded by the Ministry of Education, Science and Technology (NRF-2019R1A2C1010851). DJ is supported by the National Research Council Canada and by an NSERC Discovery Grant. LB acknowledges support from CONICYT project Basal AFB-170002. P.S. was partially supported by a Grant-in-Aid for Scientific Research (KAKENHI Number 18H01259) of Japan Society for the Promotion of Science (JSPS). Y.-L. Yang was supported, in part, by the Virginia Initiative on Cosmic Origins (VICO). JEL and HWY are supported by the Basic Science Research Program through the National Research Foundation of Korea (grant No. NRF-2018R1A2B6003423) and the Korea Astronomy and Space Science Institute under the R&D program supervised by the Ministry of Science, ICT and Future Planning.

Software: `astropy` (Astropy Collaboration et al. 2013; Price-Whelan et al. 2018), `CASA` (McMullin 2007), `XCLASS` (Möller et al. 2017)

APPENDIX

A. OBSERVED SPECTRA OF SOURCES

The spectra of all the four sources and the x-axis is the observed frequency. The red curve is the simulation result of XCLASS. Tentative detections are indicated with molecule names in brackets. The frequency bandwidth of each figure is 1 GHz for G211.47-19.27S and G208.68-19.20N1, and 2 GHz in G210.49-19.79W and G192.12-11.10.

Figure A1. Spectra of G2111.47-19.27S.

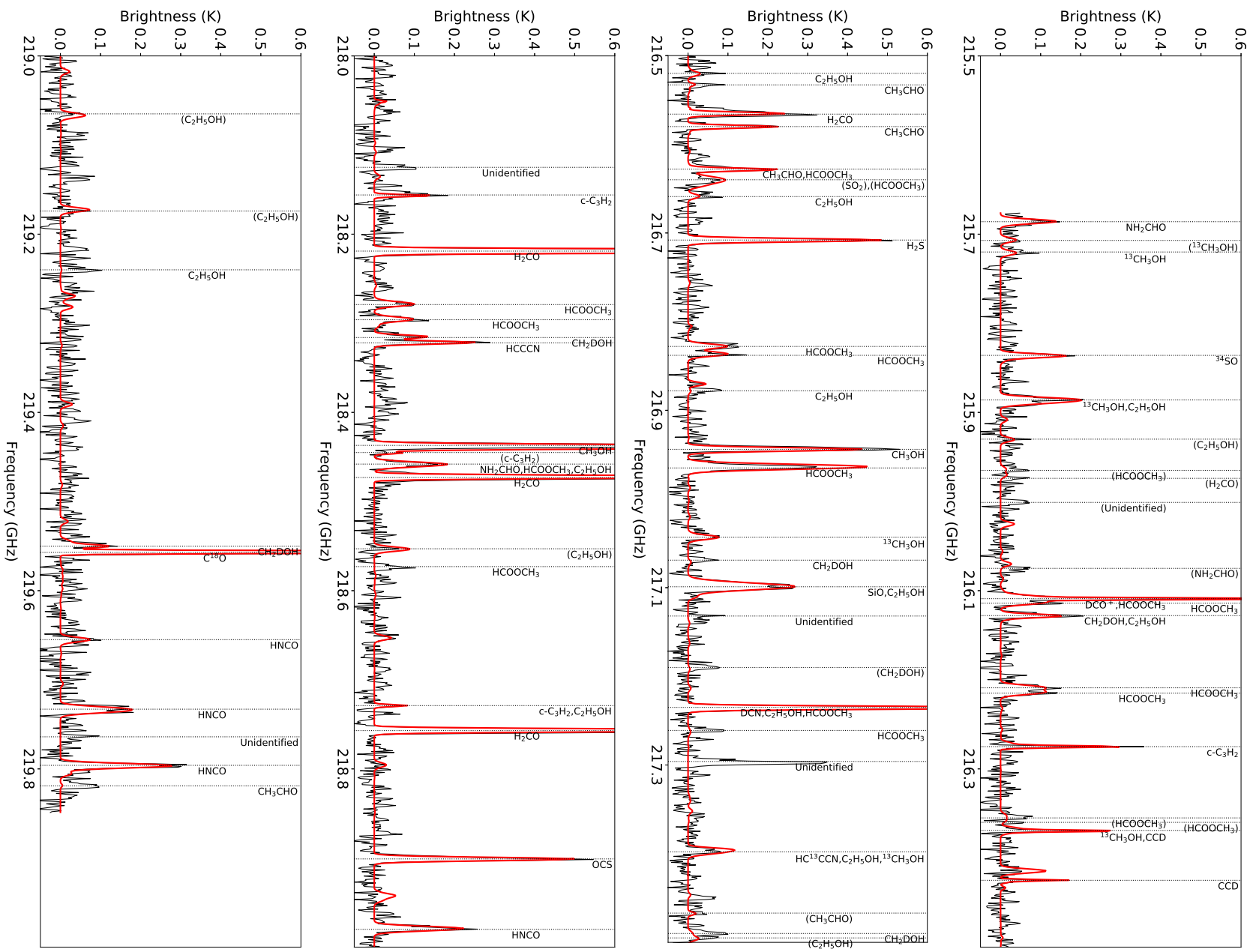


Figure A2. Spectra of G208.68-19.20N1.

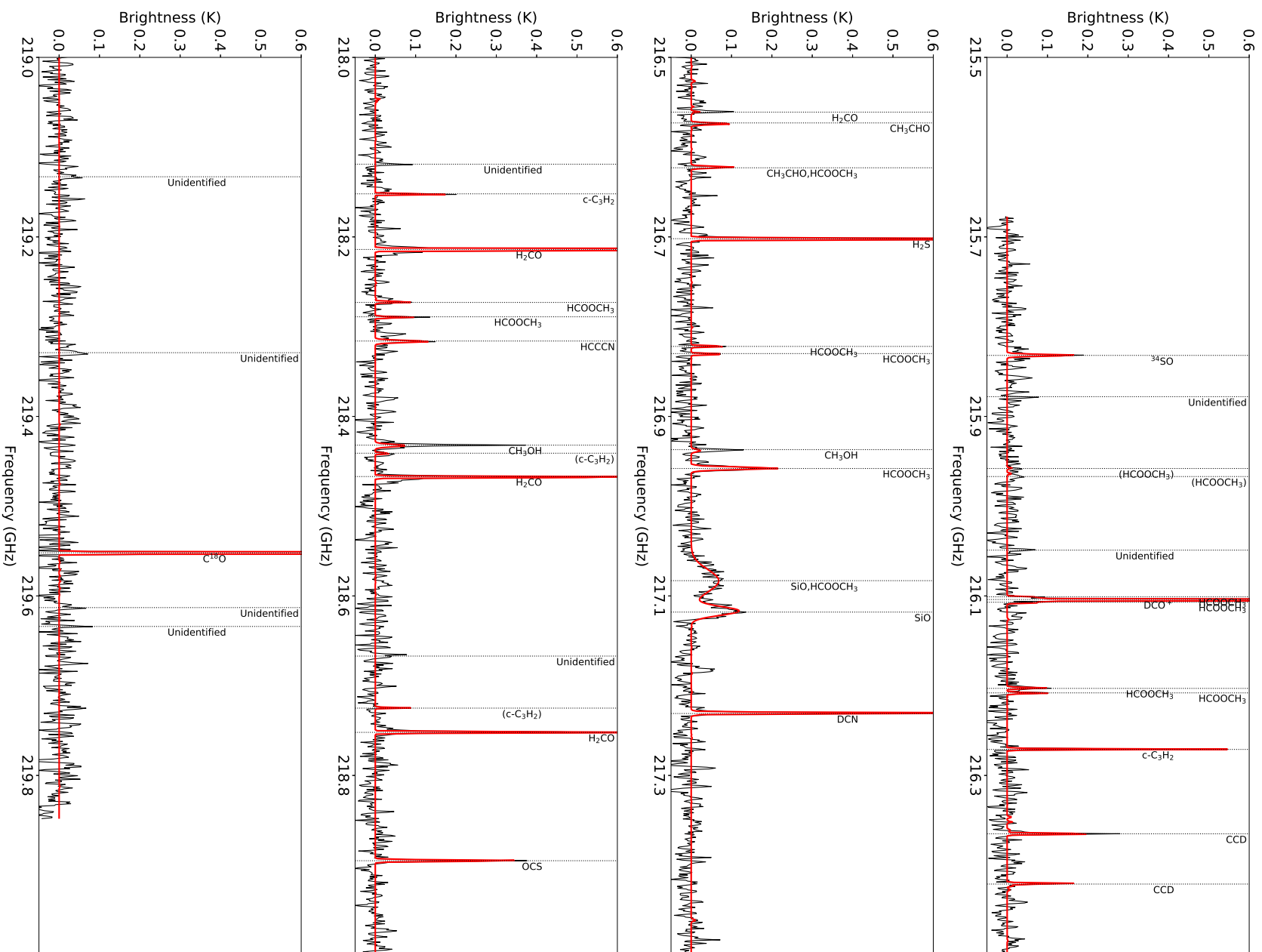


Figure A2 (continued)

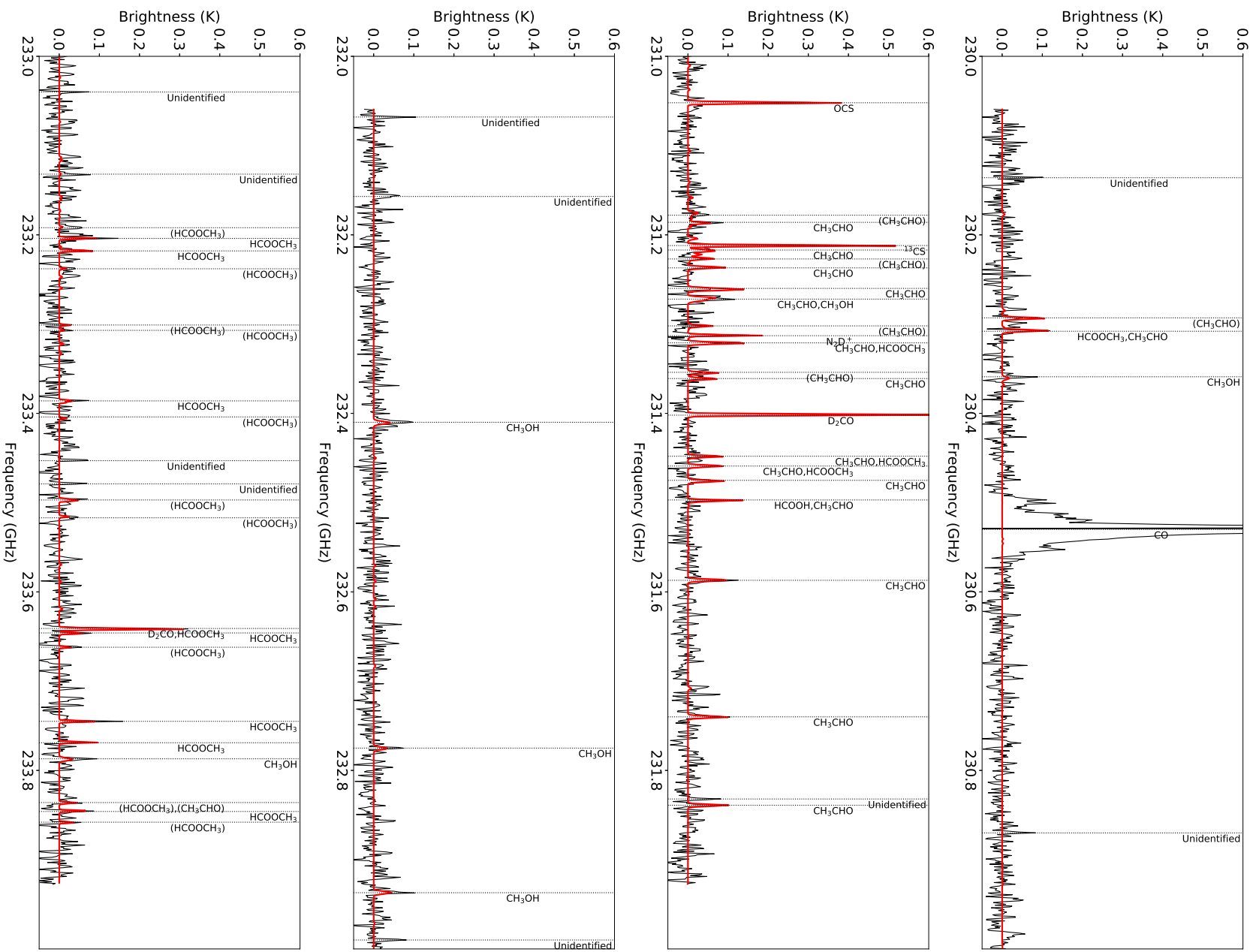


Figure A3. Spectra of G210.49-19.79W.

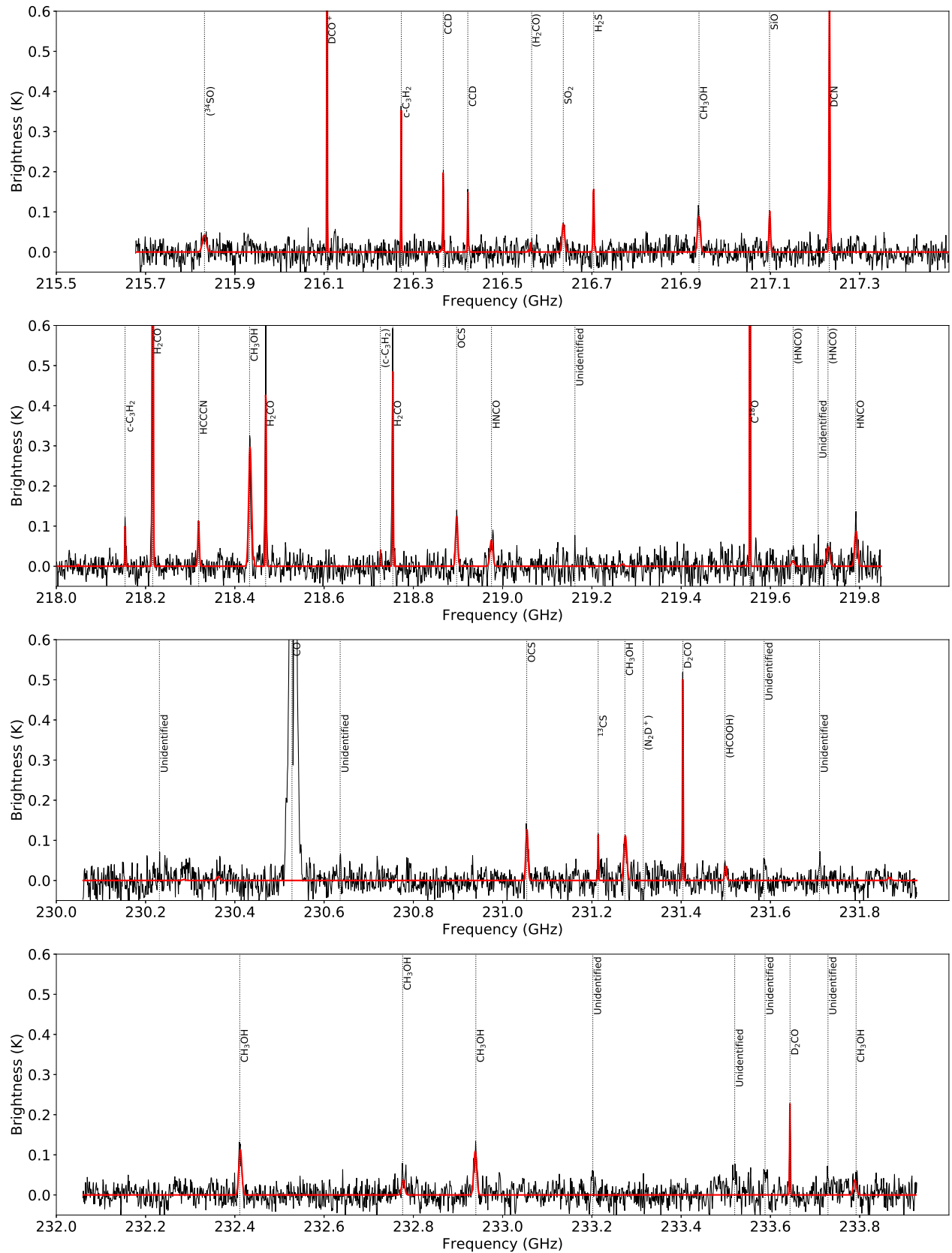
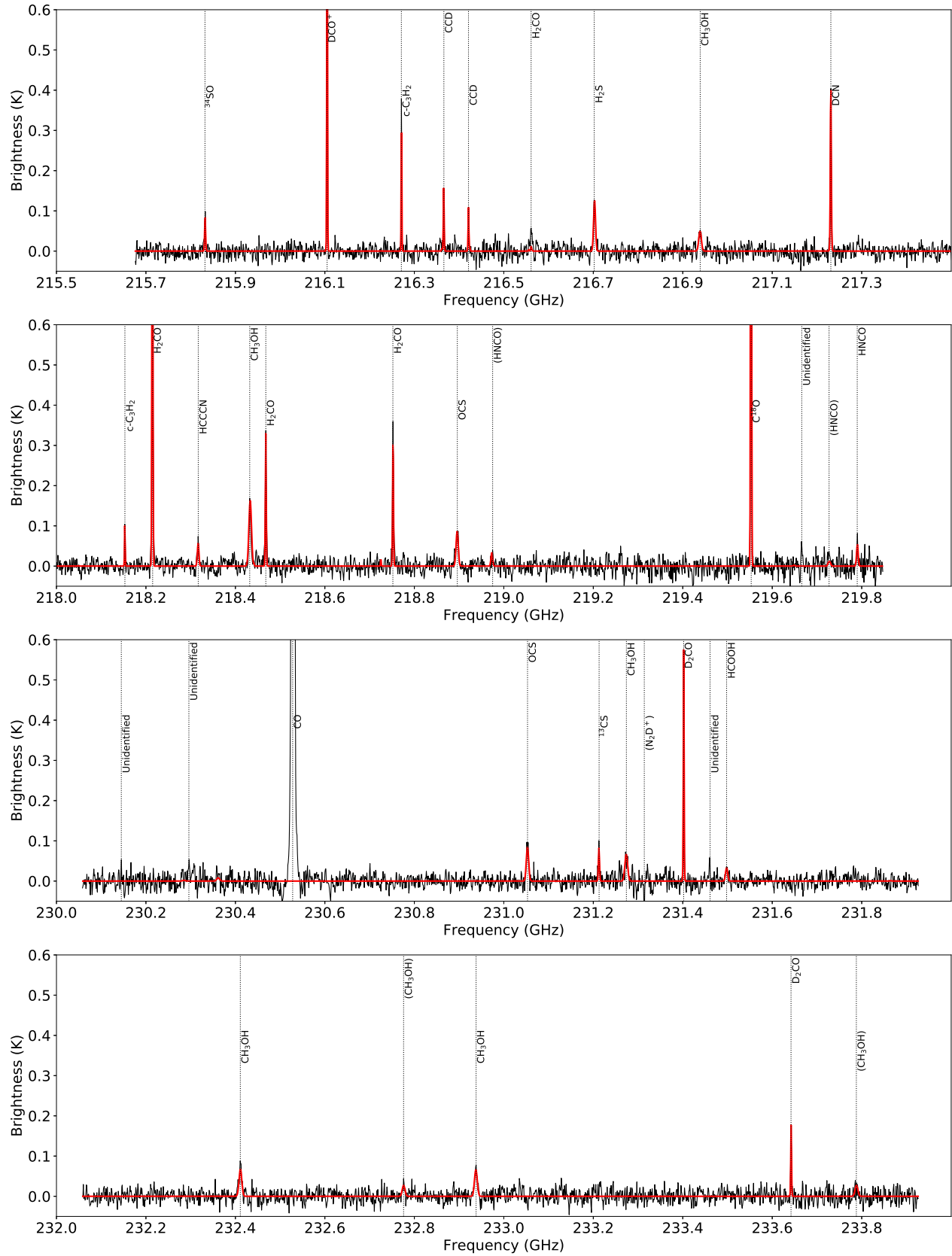


Figure A4. Spectra of G192.12-11.10.



B. MOLECULAR LINE LIST

Detected transitions of all the four sources. The f^{obs} is the observed frequency. The T_p is the peak brightness temperature. The Δv is the FWHM line width. The W is the integrated brightness within the FWHM of each line. In the Formula column, the rest frequencies in MHz are provided in bracket. It is denoted as '*N' if there are N transitions of the same species at the same rest frequency. In the Note column:

- U: Unidentified transition.
- I: Line ignored in the fitting. They are CO $J=2-1$ transitions lack of Gaussian line profile.
- N: Negative.
- T: Tentative detection. Their intensity is below $3\sigma_{\text{Chn}}$.
- B: Blended. They are different transitions of either the same species or different species.
- UE: Under-estimated by XCLASS. The observed brightness temperature is above $3\sigma_{\text{Chn}}$ but the brightness temperature calculated by XCLASS is below $2\sigma_{\text{Chn}}$. There may be blended and unidentified transitions or the other components of molecules.

Table B1. Molecular line list of G211.47-19.27S.

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
215686	0.14	6.95	0.79	NH ₂ CHO[215688]	
215707	0.05	4.17	0.18	¹³ CH ₃ OH[215708]	T
215720	0.09	4.17	0.3	¹³ CH ₃ OH[215722]	UE
215836	0.17	7.03	0.97	³⁴ SO[215840]	
215886	0.18	6.94	1.01	¹³ CH ₃ OH[215887]; C ₂ H ₅ OH[215890]	B
215930	0.07	2.78	0.16	C ₂ H ₅ OH[215932]	T
215965	0.07	2.78	0.16	HCOOCH ₃ [215966]	T
215974	0.07	2.78	0.16	H ₂ CO[215976]	T
216001	0.07	5.55	0.31	Unidentified	T, U
216075	0.07	2.29	0.13	NH ₂ CHO[216077]	T
216109	0.79	2.74	1.75	DCO ⁺ [216113]; HCOOCH ₃ [216110]	B
216114	0.16	4.16	0.54	HCOOCH ₃ [216115, 216116]	B
216128	0.21	5.2	0.88	CH ₂ DOH[216130]; C ₂ H ₅ OH[216127*2]	B
216209	0.13	8.56	0.9	HCOOCH ₃ [216211]	
216215	0.13	6.17	0.65	HCOOCH ₃ [216217]	
216275	0.36	2.23	0.65	c-C ₃ H ₂ [216279]	
216355	0.07	3.79	0.21	HCOOCH ₃ [216355]	T
216360	0.05	2.77	0.12	HCOOCH ₃ [216360]	T
216369	0.23	2.92	0.54	¹³ CH ₃ OH[216370]; CCD[216373*3]	B
216425	0.12	2.39	0.23	CCD[216428*2, 216429]	B
216520	0.11	2.77	0.24	C ₂ H ₅ OH[216522]	UE
216533	0.11	2.77	0.24	CH ₃ CHO[216534]	UE
216566	0.27	6.92	1.51	H ₂ CO[216569]	
216580	0.24	5.09	0.99	CH ₃ CHO[216582]	
216628	0.22	5.17	0.92	CH ₃ CHO[216630]; HCOOCH ₃ [216631]	B
216640	0.07	6.92	0.39	SO ₂ [216643]; HCOOCH ₃ [216638]	T, B
216659	0.09	2.77	0.2	C ₂ H ₅ OH[216660]	UE
216708	0.5	5.51	2.23	H ₂ S[216710]	
216828	0.12	6.91	0.67	HCOOCH ₃ [216830]	
216838	0.14	2.77	0.31	HCOOCH ₃ [216839]	
216878	0.08	4.15	0.27	C ₂ H ₅ OH[216879]	UE
216944	0.5	6.91	2.8	CH ₃ OH[216946]	
216965	0.33	7.06	1.89	HCOOCH ₃ [216963, 216964, 216965, 216966*2, 216967, 216968, 216969]	B
217043	0.09	6.91	0.5	¹³ CH ₃ OH[217045]	
217069	0.08	4.14	0.27	CH ₂ DOH[217072]	UE

Table B1 (continued)

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
217099	0.27	11.05	2.42	SiO[217105]; C ₂ H ₅ OH[217100]	B
217132	0.09	0.94	0.07	Unidentified	U
217190	0.07	5.52	0.32	CH ₂ DOH[217190]	T
217235	0.89	5.33	3.84	DCN[217237, 217238*2, 217239*2, 217241]; C ₂ H ₅ OH[217236, 217238*2]; HCOOCH ₃ [217236]	B
217261	0.1	5.52	0.44	HCOOCH ₃ [217263]	UE
217296	0.36	5.7	1.66	Unidentified	U
217398	0.09	4.39	0.32	HC ¹³ CCN[217399]; C ₂ H ₅ OH[217400]; ¹³ CH ₃ OH[217400]	B
217467	0.03	5.51	0.13	CH ₃ CHO[217469]	T
217490	0.11	5.51	0.48	CH ₂ DOH[217492]	UE
217495	0.07	4.14	0.22	C ₂ H ₅ OH[217497*2]	T, B
218125	0.12	4.12	0.39	Unidentified	U
218156	0.16	5.94	0.77	c-C ₃ H ₂ [218160]	
218219	2.58	4.49	9.38	H ₂ CO[218222]	
218279	0.11	5.49	0.49	HCOOCH ₃ [218281]	
218296	0.12	5.49	0.53	HCOOCH ₃ [218298]	
218316	0.11	5.97	0.53	CH ₂ DOH[218316]	
218322	0.29	4.24	1.0	HCCCN[218325]	
218437	0.94	5.9	4.49	CH ₃ OH[218440]	
218445	0.05	2.74	0.11	c-C ₃ H ₂ [218449*2]	T, B
218458	0.16	6.31	0.82	NH ₂ CHO[218460]; HCOOCH ₃ [218458]; C ₂ H ₅ OH[218461]	B
218473	1.19	5.49	5.29	H ₂ CO[218476]	
218553	0.07	2.69	0.15	C ₂ H ₅ OH[218554]	T
218573	0.1	4.11	0.33	HCOOCH ₃ [218575]	UE
218729	0.08	3.53	0.23	c-C ₃ H ₂ [218733*2]; C ₂ H ₅ OH[218729]	B
218757	1.15	5.77	5.37	H ₂ CO[218760]	
218901	0.52	5.37	2.26	OCS[218903]	
218980	0.24	5.89	1.15	HNCO[218981]	
219065	0.05	2.74	0.11	C ₂ H ₅ OH[219067, 219069]	T, B
219174	0.07	2.74	0.16	C ₂ H ₅ OH[219174]	T
219240	0.1	4.1	0.33	C ₂ H ₅ OH[219241*2]	B, UE
219550	0.18	2.73	0.4	CH ₂ DOH[219551]	
219557	1.84	3.42	5.1	C ¹⁸ O[219560]	
219655	0.09	5.82	0.42	HNCO[219657*2]	B
219733	0.18	10.91	1.59	HNCO[219734, 219737]	B
219764	0.09	2.73	0.2	Unidentified	U
219796	0.27	8.18	1.79	HNCO[219798]	
219819	0.12	5.46	0.52	CH ₃ CHO[219820]	UE
230104	0.13	5.21	0.53	CH ₃ CHO[230106]	UE
230141	0.07	1.6	0.09	C ₂ H ₅ OH[230145]	
230229	0.09	6.51	0.47	C ₂ H ₅ OH[230231]	
230253	0.15	4.11	0.5	CH ₂ DOH[230254]	
230300	0.24	5.35	1.04	CH ₃ CHO[230302]	
230314	0.31	4.96	1.25	CH ₃ CHO[230316]; HCOOCH ₃ [230316]	B
230367	0.12	6.38	0.62	CH ₃ OH[230369]	
230375	0.09	3.85	0.28	HCOOCH ₃ [230377]; CH ₂ DOH[230377]	B, UE
230537	CO[230538]	I
230672	0.11	3.27	0.29	C ₂ H ₅ OH[230673]	
230793	0.08	1.89	0.12	C ₂ H ₅ OH[230794*2]	B
230952	0.07	5.19	0.29	C ₂ H ₅ OH[230954]	
230965	0.09	2.6	0.19	Unidentified	U
230990	0.08	3.89	0.25	C ₂ H ₅ OH[230991]	
231059	0.52	5.11	2.15	OCS[231061]	
231077	0.1	0.85	0.07	Unidentified	U

Table B1 (continued)

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
231149	0.07	2.59	0.15	Unidentified	U
231194	0.09	2.74	0.2	CH ₃ CHO[231195, 231196]	B
231217	0.32	3.89	1.01	¹³ CS[231221]; C ₂ H ₅ OH[231221]	B
231225	0.15	3.51	0.43	CH ₃ CHO[231226, 231227*2]	B
231228	0.06	6.38	0.31	CH ₃ CHO[231231]; HCOOCH ₃ [231232]	T, B
231233	0.11	4.36	0.39	CH ₃ CHO[231235*2]	B
231243	0.17	4.01	0.55	CH ₃ CHO[231245*2]	B
231268	0.26	5.05	1.06	CH ₃ CHO[231268, 231270*2]	B
231276	0.21	3.55	0.6	C ₂ H ₅ OH[231277*2]	B
231279	0.52	6.58	2.77	CH ₃ OH[231281]; CH ₃ CHO[231279]	B
231309	0.11	5.18	0.46	CH ₃ CHO[231310]	
231328	0.25	3.78	0.77	HCOOCH ₃ [231331]; CH ₃ CHO[231330*2]	B
231361	0.16	4.24	0.55	CH ₃ CHO[231363]	
231368	0.12	5.18	0.5	CH ₃ CHO[231370]	
231407	0.41	3.89	1.29	D ₂ CO[231410]	
231455	0.14	5.18	0.59	CH ₃ CHO[231457]; HCOOCH ₃ [231457]	B
231466	0.16	4.61	0.6	CH ₃ CHO[231468]; HCOOCH ₃ [231470]	B
231483	0.18	4.01	0.58	CH ₃ CHO[231484]	
231505	0.21	3.66	0.62	HCOOH[231506]; CH ₃ CHO[231506]	B
231558	0.09	6.61	0.48	C ₂ H ₅ OH[231559, 231561]; HCOOCH ₃ [231560, 231561]	B
231593	0.19	4.66	0.72	CH ₃ CHO[231595]	
231667	0.08	3.88	0.25	C ₂ H ₅ OH[231669]	
231723	0.08	1.69	0.11	Unidentified	U
231736	0.07	3.37	0.19	C ₂ H ₅ OH[231738]	
231747	0.23	5.17	0.96	CH ₃ CHO[231749]	
231839	0.11	3.81	0.34	CH ₂ DOH[231840, 231841]	B
231845	0.19	5.18	0.8	CH ₃ CHO[231848]	
231873	0.06	2.63	0.13	HNCO[231873]; C ₂ H ₅ OH[231874*2]	T, B
232075	0.15	7.82	0.95	C ₂ H ₅ OH[232076]; ¹³ CH ₃ CN[232077*2]	B
232109	0.09	5.17	0.38	Unidentified	U
232164	0.12	5.17	0.49	¹³ CH ₃ CN[232164]	UE
232177	0.07	6.46	0.38	Unidentified	U
232193	0.07	7.75	0.44	¹³ CH ₃ CN[232195*2]	B
232229	0.09	3.7	0.27	¹³ CH ₃ CN[232230]	
232233	0.09	2.58	0.19	¹³ CH ₃ CN[232234]	
232272	0.18	6.45	0.94	NH ₂ CHO[232274]	
232286	0.08	3.41	0.22	C ₂ H ₅ OH[232287]	UE
232317	0.07	2.09	0.12	C ₂ H ₅ OH[232318]	
232333	0.09	3.87	0.28	Unidentified	U
232348	0.09	3.87	0.28	Unidentified	U
232365	0.08	1.29	0.08	C ₂ H ₅ OH[232367]	UE
232403	0.11	3.87	0.34	C ₂ H ₅ OH[232405]	
232417	0.57	6.04	2.79	CH ₃ OH[232419]; C ₂ H ₅ OH[232423]	B
232427	0.1	1.29	0.11	Unidentified	U
232489	0.09	5.16	0.38	C ₂ H ₅ OH[232491]	
232575	0.08	3.87	0.24	HCOOCH ₃ [232579]	UE
232596	0.11	4.38	0.39	C ₂ H ₅ OH[232597]; HCOOCH ₃ [232597]	B
232706	0.09	3.86	0.28	Unidentified	U
232782	0.31	6.0	1.51	CH ₃ OH[232784]	
232797	0.09	3.86	0.28	NH ₂ CHO[232798]	UE
232808	0.11	5.67	0.51	C ₂ H ₅ OH[232809]	
232928	0.09	4.33	0.32	C ₂ H ₅ OH[232928]; CH ₂ DOH[232929]	B
232943	0.59	5.15	2.46	CH ₃ OH[232945]	

Table B1 (continued)

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
233030	0.13	3.86	0.39	HCOOCH ₃ [233033]	UE
233048	0.09	3.86	0.28	Unidentified	U
233081	0.15	3.86	0.48	CH ₂ DOH[233083]	UE
233132	0.19	2.57	0.39	CH ₂ DOH[233134]	UE
233141	0.09	3.73	0.27	HCOOCH ₃ [233141]; CH ₂ DOH[233142]	B
233208	0.22	7.96	1.42	CH ₂ DOH[233210]; C ₂ H ₅ OH[233209]; HCOOCH ₃ [233213*2]	B
233212	0.18	3.86	0.56	C ₂ H ₅ OH[233215]	
233225	0.12	5.14	0.5	HCOOCH ₃ [233227]; C ₂ H ₅ OH[233229*2]	B
233246	0.06	2.57	0.12	HCOOCH ₃ [233247*2]	T, B
233308	0.08	5.14	0.33	HCOOCH ₃ [233310, 233310]	B, UE
233394	0.12	5.14	0.5	HCOOCH ₃ [233395*2, 233397]	B
233414	0.05	2.57	0.1	HCOOCH ₃ [233414]	T
233437	0.09	1.93	0.14	Unidentified	U
233441	0.09	2.57	0.19	CH ₃ CHO[233443]	UE
233459	0.12	6.42	0.62	CH ₂ DOH[233461]	
233488	0.15	3.85	0.48	NH ₂ CHO[233490*2, 233492*2]	B, UE
233493	0.12	3.85	0.37	HCOOH[233493*2]	B, UE
233497	0.16	3.85	0.51	NH ₂ CHO[233499*2]	B, UE
233505	0.19	4.77	0.73	NH ₂ CHO[233504*2, 233505*2]; HCOOCH ₃ [233505*2, 233507*2]	B
233527	0.16	5.44	0.71	NH ₂ CHO[233529*2]; HCOOCH ₃ [233525]	B
233570	0.11	2.97	0.26	C ₂ H ₅ OH[233571]	
233594	0.19	4.98	0.77	NH ₂ CHO[233595, 233596]; HCOOCH ₃ [233598]	B
233598	0.09	2.57	0.19	C ₂ H ₅ OH[233602]	
233648	0.23	6.42	1.2	D ₂ CO[233650]; HCOOCH ₃ [233650]	B
233653	0.08	4.11	0.27	HCOOCH ₃ [233655*2]	B
233669	0.09	2.12	0.15	HCOOCH ₃ [233671]	UE
233734	0.16	4.81	0.62	CH ₂ DOH[233736]; NH ₂ CHO[233736]	B
233745	0.16	4.81	0.62	NH ₂ CHO[233747]	
233752	0.14	3.85	0.44	HCOOCH ₃ [233754]	
233777	0.12	3.85	0.37	HCOOCH ₃ [233778]	
233794	0.39	4.82	1.52	CH ₃ OH[233796]	
233844	0.1	4.75	0.38	HCOOCH ₃ [233845]; CH ₃ CHO[233845]	B
233853	0.09	3.85	0.28	HCOOCH ₃ [233854*2]	B
233866	0.04	2.65	0.09	HCOOCH ₃ [233867]	T
233895	0.17	5.87	0.81	NH ₂ CHO[233897]	

Table B2. Molecular line list of G208.68-19.20N1.

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
215832	0.19	2.63	0.4	³⁴ SO[215840]	
215878	0.08	1.39	0.09	Unidentified	U
215958	0.04	3.07	0.1	HCOOCH ₃ [215966]	T
215967	0.05	4.16	0.15	HCOOCH ₃ [215976]	T
216049	0.07	2.78	0.16	Unidentified	U
216101	0.07	2.77	0.16	HCOOCH ₃ [216110]	
216104	1.72	2.01	2.81	DCO ⁺ [216113]	
216107	0.07	1.17	0.09	HCOOCH ₃ [216116*2]	B
216203	0.1	4.33	0.34	HCOOCH ₃ [216211]	
216208	0.12	2.0	0.2	HCOOCH ₃ [216217]	
216271	0.54	1.39	0.61	c-C ₃ H ₂ [216279]	
216365	0.28	2.15	0.49	CCD[216373*3]	B

Table B2 (continued)

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
216421	0.12	2.17	0.21	CCD[216428*2, 216429, 216430, 216430, 216431]	B
216561	0.11	2.68	0.24	H ₂ CO[216569]	UE
216573	0.09	1.38	0.1	CH ₃ CHO[216582]	
216623	0.09	2.77	0.2	CH ₃ CHO[216630]; HCOOCH ₃ [216631]	B
216702	1.08	2.77	2.42	H ₂ S[216710]	
216822	0.09	1.64	0.12	HCOOCH ₃ [216830]	
216830	0.09	2.77	0.2	HCOOCH ₃ [216839]	
216937	0.14	4.15	0.45	CH ₃ OH[216946]	
216958	0.16	7.36	0.93	HCOOCH ₃ [216963, 216964, 216965, 216966*2, 216967, 216968, 216969]	B
217083	0.09	21.95	1.6	SiO[217105]; HCOOCH ₃ [217077]	B
217118	0.13	13.01	1.34	SiO[217105]	
217231	0.9	2.07	1.51	DCN[217237, 217238, 217239*3, 217241]	B
218119	0.09	1.37	0.1	Unidentified	U
218152	0.23	2.06	0.38	c-C ₃ H ₂ [218160]	
218214	3.47	1.93	5.43	H ₂ CO[218222]	
218273	0.07	5.41	0.33	HCOOCH ₃ [218281]	
218289	0.14	2.06	0.23	HCOOCH ₃ [218298]	
218316	0.14	4.12	0.48	HCCCN[218325]	
218432	0.36	2.62	0.76	CH ₃ OH[218440]	
218441	0.05	4.12	0.15	c-C ₃ H ₂ [218449*2]	T, B
218467	0.64	2.72	1.41	H ₂ CO[218476]	
218667	0.07	2.74	0.16	Unidentified	U
218725	0.06	1.37	0.07	c-C ₃ H ₂ [218733*2]	T, B
218752	0.64	2.41	1.25	H ₂ CO[218760]	
218895	0.37	2.26	0.68	OCS[218903]	
219133	0.07	2.74	0.16	Unidentified	U
219329	0.07	4.1	0.24	Unidentified	U
219552	5.25	1.73	7.51	C ¹⁸ O[219560]	
219613	0.07	2.73	0.16	Unidentified	U
219634	0.07	2.05	0.12	Unidentified	U
230136	0.1	1.95	0.16	Unidentified	U
230293	0.05	3.14	0.13	CH ₃ CHO[230302]	T
230308	0.12	1.92	0.19	HCOOCH ₃ [230316]; CH ₃ CHO[230316]	B
230359	0.09	2.23	0.17	CH ₃ OH[230369]	UE
230530	CO[230538]	I
230870	0.08	1.95	0.13	Unidentified	U
231052	0.39	2.34	0.75	OCS[231061]	
231178	0.05	2.59	0.11	CH ₃ CHO[231184]	T
231186	0.09	2.7	0.2	CH ₃ CHO[231195*2]	B
231212	0.57	1.93	0.89	¹³ CS[231221]	
231217	0.06	5.31	0.26	CH ₃ CHO[231226, 231227*2]	B
231227	0.05	2.49	0.09	CH ₃ CHO[231235*2]	T, B
231237	0.08	2.47	0.15	CH ₃ CHO[231245*2]	B
231260	0.06	4.25	0.22	CH ₃ CHO[231268, 231270*2]	B
231272	0.11	4.44	0.4	CH ₃ CHO[231279]; CH ₃ OH[231281]	B
231302	0.05	2.59	0.09	CH ₃ CHO[231310*2]	T, B
231313	0.18	2.15	0.3	N ₂ D ⁺ [231320*3, 231321*4, 231322]	B
231321	0.09	3.43	0.24	CH ₃ CHO[231330*2]; HCOOCH ₃ [231331]	B
231354	0.03	6.07	0.14	CH ₃ CHO[231363]	T
231361	0.06	2.74	0.13	CH ₃ CHO[231370]	
231402	0.71	2.0	1.15	D ₂ CO[231410]	
231448	0.09	1.9	0.13	CH ₃ CHO[231457]; HCOOCH ₃ [231457]	B
231459	0.09	2.77	0.2	CH ₃ CHO[231468]; HCOOCH ₃ [231470]	B

Table B2 (continued)

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
231475	0.07	3.58	0.2	CH ₃ CHO[231484]	
231497	0.09	2.59	0.19	HCOOH[231506]; CH ₃ CHO[231506]	B
231587	0.12	3.16	0.32	CH ₃ CHO[231595]	
231740	0.1	3.35	0.26	CH ₃ CHO[231749]	
231832	0.08	2.59	0.17	Unidentified	U
231839	0.1	1.96	0.15	CH ₃ CHO[231848]	
232068	0.11	0.86	0.07	Unidentified	U
232157	0.07	2.58	0.15	Unidentified	U
232410	0.1	4.23	0.34	CH ₃ OH[232419]	
232775	0.09	2.23	0.15	CH ₃ OH[232784]	
232937	0.09	3.19	0.24	CH ₃ OH[232945]	
232990	0.08	2.57	0.17	Unidentified	U
233040	0.08	1.93	0.13	Unidentified	U
233132	0.08	1.93	0.13	Unidentified	U
233192	0.06	3.65	0.19	HCOOCH ₃ [233200*2]	T, B
233204	0.15	2.12	0.25	HCOOCH ₃ [233213*2]	B
233218	0.08	2.25	0.15	HCOOCH ₃ [233227]	
233238	0.04	3.22	0.11	HCOOCH ₃ [233247*2]	T, B
233301	0.05	1.29	0.05	HCOOCH ₃ [233310*2]	T, B
233307	0.05	1.53	0.05	HCOOCH ₃ [233316]	T
233386	0.09	1.57	0.11	HCOOCH ₃ [233395*2, 233397]	B, U
233404	0.04	2.57	0.07	HCOOCH ₃ [233414]	T
233453	0.08	1.93	0.13	Unidentified	U
233479	0.08	1.93	0.13	Unidentified	U
233497	0.06	2.93	0.13	HCOOCH ₃ [233505, 233507*2]	T, B
233517	0.04	6.5	0.21	HCOOCH ₃ [233525]	T
233641	0.32	2.24	0.59	D ₂ CO[233650]; HCOOCH ₃ [233650]	B
233646	0.09	2.36	0.16	HCOOCH ₃ [233655*2]	B
233662	0.06	2.31	0.12	HCOOCH ₃ [233671]	T
233745	0.16	1.97	0.26	HCOOCH ₃ [233754]	
233769	0.07	1.92	0.11	HCOOCH ₃ [233778]	
233787	0.09	3.34	0.24	CH ₃ OH[233796]	
233836	0.06	1.54	0.08	HCOOCH ₃ [233845]; CH ₃ CHO[233845]	T, B
233846	0.09	1.85	0.13	HCOOCH ₃ [233854*2]	B
233858	0.06	1.92	0.1	HCOOCH ₃ [233867]	T

Table B3. Molecular line list of G210.49-19.79W.

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
215832	0.05	11.11	0.43	³⁴ SO[215840]	T
216107	1.12	2.77	2.52	DCO ⁺ [216113]	
216273	0.32	2.77	0.72	c-C ₃ H ₂ [216279]	
216367	0.23	2.25	0.41	CCD[216373*3]	B
216422	0.18	2.05	0.29	CCD[216428*2, 216429]	B
216565	0.03	4.24	0.11	H ₂ CO[216569]	T
216636	0.07	6.92	0.41	SO ₂ [216643]	
216704	0.16	4.31	0.55	H ₂ S[216710]	
216940	0.1	8.09	0.91	CH ₃ OH[216946]	
217098	0.1	4.33	0.48	SiO[217105]	
217232	0.72	3.18	1.86	DCN[217237, 217238, 217239*3, 217241]	B
218154	0.12	3.57	0.33	c-C ₃ H ₂ [218160]	

Table B3 (continued)

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
218216	2.15	2.77	4.83	H ₂ CO[218222]	
218319	0.11	4.78	0.43	HCCCN[218325]	
218433	0.32	5.46	1.42	CH ₃ OH[218440]	
218469	0.56	3.86	1.74	H ₂ CO[218476]	
218726	0.04	4.11	0.13	c-C ₃ H ₂ [218733*2]	T, B
218754	0.58	3.69	1.73	H ₂ CO[218760]	
218897	0.12	6.85	0.67	OCS[218903]	
218975	0.08	10.95	0.71	HNCO[218981]	
219162	0.08	2.05	0.13	Unidentified	U
219554	3.15	2.73	6.94	C ¹⁸ O[219560]	
219651	0.05	5.46	0.24	HNCO[219657*2]	T, B
219707	0.08	2.05	0.13	Unidentified	U
219729	0.05	5.46	0.24	HNCO[219734*2]	T, B
219791	0.13	5.31	0.54	HNCO[219798]	
230231	0.07	1.3	0.08	Unidentified	U
230528	CO[230538]	I
230636	0.06	3.25	0.17	Unidentified	U
231054	0.13	5.88	0.78	OCS[231061]	
231214	0.12	2.05	0.2	¹³ CS[231221]	
231274	0.11	6.48	0.57	CH ₃ OH[231281]	
231315	-0.17	2.59	-0.35	N ₂ D ⁺ [231320]	N
231404	0.55	2.15	0.96	D ₂ CO[231410]	
231498	0.04	6.01	0.21	HCOOH[231506]	T
231586	0.06	4.51	0.23	Unidentified	U
231710	0.07	5.18	0.28	Unidentified	U
232411	0.12	7.58	0.97	CH ₃ OH[232419]	
232776	0.09	3.86	0.28	CH ₃ OH[232784]	
232940	0.1	10.14	1.08	CH ₃ OH[232945]	
233202	0.07	5.14	0.28	Unidentified	U
233520	0.08	6.42	0.42	Unidentified	U
233588	0.07	6.42	0.37	Unidentified	U
233644	0.22	2.87	0.51	D ₂ CO[233650]	
233729	0.07	3.21	0.18	Unidentified	U
233792	0.07	3.85	0.22	CH ₃ OH[233796]	UE

Table B4. Molecular line list of G192.12-11.10.

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
215832	0.09	2.77	0.21	³⁴ SO[215840]	
216105	1.37	2.3	2.55	DCO ⁺ [216113]	
216271	0.39	2.1	0.66	c-C ₃ H ₂ [216279]	
216366	0.15	2.87	0.36	CCD[216373*3]	B
216421	0.11	2.51	0.23	CCD[216428*2, 216429]	B
216561	0.06	6.48	0.29	H ₂ CO[216569]	UE
216702	0.14	6.92	0.76	H ₂ S[216710]	
216939	0.05	6.91	0.27	CH ₃ OH[216946]	
217231	0.41	3.47	1.16	DCN[217237, 217238, 217239*3]	B
218153	0.11	2.32	0.21	c-C ₃ H ₂ [218160]	
218215	1.55	3.03	3.8	H ₂ CO[218222]	
218317	0.06	4.12	0.21	HCCCN[218325]	
218432	0.16	6.01	0.76	CH ₃ OH[218440]	

Table B4 (*continued*)

f^{obs} MHz	T_p K	Δv km/s	W K km/s	Formula [f^{rest}] MHz	Note
218468	0.33	3.94	1.04	H ₂ CO[218476]	
218752	0.34	3.82	1.06	H ₂ CO[218760]	
218896	0.08	8.22	0.53	OCS[218903]	
218975	0.04	2.74	0.09	HNCO[218981]	T
219553	3.68	2.62	7.8	C ¹⁸ O[219560]	
219666	0.05	4.09	0.18	Unidentified	U
219727	0.03	7.11	0.2	HNCO[219737]	T
219790	0.08	3.2	0.2	HNCO[219798]	
230145	0.05	1.95	0.09	Unidentified	U
230296	0.05	3.91	0.17	Unidentified	U
230528	CO[230538]	I
231053	0.09	6.11	0.55	OCS[231061]	
231213	0.09	2.75	0.21	¹³ CS[231221]	
231274	0.07	7.13	0.42	CH ₃ OH[231281]	
231314	-0.11	1.61	-0.18	N ₂ D ⁺ [231320]	N
231402	0.57	2.62	1.21	D ₂ CO[231410]	
232411	0.1	6.45	0.5	CH ₃ OH[232419]	
231461	0.05	3.89	0.17	Unidentified	U
231498	0.04	6.13	0.18	HCOOH[231506]	T
232776	0.04	3.52	0.12	CH ₃ OH[232784]	T
232938	0.07	6.51	0.48	CH ₃ OH[232945]	
233642	0.17	2.86	0.4	D ₂ CO[233650]	
233788	0.02	12.08	0.2	CH ₃ OH[233796]	T

C. TRANSITION LIST

Table C1. List of identified transitions.

Name	Formula	f_{rest} (MHz)	Transition Number	E_u (K)	g_u	A_u (s^{-1})	$I_{\text{CDMS.JPL}}$	Ref.	
Acetaldehyde	CH ₃ CHO	216534	14(3,11) - 14(2,12) E	117.68	58	3.45e-05	-4.3787	JPL	
	CH ₃ CHO	216582	11(1,10) - 10(1,9) E	64.87	46	3.55e-04	-3.3908	JPL	
	CH ₃ CHO	216630	11(1,10) - 10(1,9) A	64.81	46	3.55e-04	-3.3908	JPL	
	CH ₃ CHO	217469	14(3,11)-14(2,12) A	117.71	58	3.82e-05	-4.3365	JPL	
	CH ₃ CHO	219820	4(2,3) - 3(1,3) E	1176.53	63	1.28e-07	-5.1	JPL	
	CH ₃ CHO	230106	26(3,23) - 25(4,22) E	348.37	106	1.41e-05	-4.8651	JPL	
	CH ₃ CHO	230302	12(2,11) - 11(2,10) A	81.04	50	4.19e-04	-3.3313	JPL	
	CH ₃ CHO	230316	12(2,11) - 11(2,10) E	81.05	50	4.19e-04	-3.3314	JPL	
	CH ₃ CHO	231195	12(8,4) - 11(8,3) E	216.41	50	2.43e-04	-3.7663	JPL	
	CH ₃ CHO	231196	12(9,4) - 11(9,3) E	254.62	50	1.91e-04	-3.9254	JPL	
	CH ₃ CHO	231226	12(7,5) - 11(7,4) E	182.65	50	2.88e-04	-3.6428	JPL	
	CH ₃ CHO	231227	12(9,3) - 11(9,2) A	254.53	50	1.91e-04	-3.9253	JPL	
	CH ₃ CHO	231227	12(9,4) - 11(9,3) A	254.53	50	1.91e-04	-3.9253	JPL	
	CH ₃ CHO	231231	12(8,5) - 11(8,4) E	216.33	50	2.43e-04	-3.7662	JPL	
	CH ₃ CHO	231235	12(8,4) - 11(8,3) A	216.30	50	2.43e-04	-3.7662	JPL	
	CH ₃ CHO	231235	12(8,5) - 11(8,4) A	216.30	50	2.43e-04	-3.7662	JPL	
	CH ₃ CHO	231245	12(7,5) - 11(7,4) A	182.58	50	2.88e-04	-3.6426	JPL	
	CH ₃ CHO	231245	12(7,6) - 11(7,5) A	182.58	50	2.88e-04	-3.6426	JPL	
	CH ₃ CHO	231268	12(7,6) - 11(7,5) E	182.54	50	2.88e-04	-3.6426	JPL	
	CH ₃ CHO	231270	12(6,7) - 11(6,6) A	153.36	50	3.28e-04	-3.5445	JPL	
	CH ₃ CHO	231270	12(6,6) - 11(6,5) A	153.36	50	3.28e-04	-3.5445	JPL	
	CH ₃ CHO	231279	12(6,6) - 11(6,5) E	153.35	50	3.28e-04	-3.5446	JPL	
	CH ₃ CHO	231310	12(6,7) - 11(6,6) E	153.26	50	3.28e-04	-3.5444	JPL	
	CH ₃ CHO	231330	12(5,8) - 11(5,7) A	128.61	50	3.62e-04	-3.4663	JPL	
	CH ₃ CHO	231330	12(5,7) - 11(5,6) A	128.61	50	3.62e-04	-3.4663	JPL	
	CH ₃ CHO	231363	12(5,7) - 11(5,6) E	128.54	50	3.62e-04	-3.4663	JPL	
	CH ₃ CHO	231370	12(5,8) - 11(5,7) E	128.51	50	3.62e-04	-3.4662	JPL	
	CH ₃ CHO	231457	12(4,9) - 11(4,8) A	108.35	50	3.90e-04	-3.4049	JPL	
	CH ₃ CHO	231468	12(4,8) - 11(4,7) A	108.36	50	3.90e-04	-3.4049	JPL	
	CH ₃ CHO	231484	12(4,8) - 11(4,7) E	108.29	50	3.90e-04	-3.4048	JPL	
	CH ₃ CHO	231506	12(4,9) - 11(4,8) E	108.25	50	3.90e-04	-3.4048	JPL	
	CH ₃ CHO	231595	12(3,10) - 11(3,9) A	92.57	50	4.12e-04	-3.3586	JPL	
	CH ₃ CHO	231749	12(3,10) - 11(3,9) E	92.51	50	4.09e-04	-3.3615	JPL	
	CH ₃ CHO	231848	12(3,9) - 11(3,8) E	92.61	50	4.10e-04	-3.3611	JPL	
	CH ₃ CHO	233443	17(2,16) - 17(0,17) A	149.96	70	2.35e-06	-5.5424	JPL	
	CH ₃ CHO	233845	6(3,3) - 6(2,4) E	39.81	26	3.51e-05	-4.6399	JPL	
	Carbon Monosulfide (¹³ C-substituted)	¹³ CS	231221	5 - 4	33.29	22	2.51e-04	-1.6401	CDMS
	Carbon Monoxide	CO	230538	2 - 1	16.59	5	6.91e-07	-4.1197	CDMS
	Carbon Monoxide (¹⁸ O-substituted)	C ¹⁸ O	219560	2 - 1	15.81	5	6.01e-07	-4.1794	CDMS
	Carbonyl Sulfide	OCS	218903	3(2,1) - 2(2,0)	99.81	37	3.04e-05	-3.0903	CDMS
	OCS	231061	19 - 18	110.90	39	3.58e-05	-2.6263	CDMS	
Cyanoacetylene	HCCCN	218325	24 - 23	130.98	49	8.26e-04	-1.2947	CDMS	
Cyanoacetylene (¹³ C-substituted)	HC ¹³ CCN	217399	24 - 23	130.42	49	8.16e-04	-1.2994	CDMS	
Cyclopropenylidene	c-C ₃ H ₂	216279	3(3,0) - 2(2,1)	19.47	21	2.56e-04	-2.9741	CDMS	
	c-C ₃ H ₂	218160	5(2,4) - 4(1,3)	35.42	11	4.04e-04	-3.0831	CDMS	
	c-C ₃ H ₂	218449	8(2,6) - 8(1,7)	86.93	17	1.49e-04	-3.4013	CDMS	
	c-C ₃ H ₂	218449	8(3,6) - 8(2,7)	86.93	51	1.49e-04	-2.9242	CDMS	
	c-C ₃ H ₂	218733	7(1,6) - 7(0,7)	61.17	15	8.93e-05	-3.1652	CDMS	
	c-C ₃ H ₂	218733	7(2,6) - 7(1,7)	61.17	45	8.93e-05	-3.6423	CDMS	
Diazenylium (Deuterated)	N ₂ D ⁺	231320	3(3,2) - 2(3,2)	22.20	5	8.82e-05	-3.1516	CDMS	
	N ₂ D ⁺	231320	3(3,4) - 2(3,4)	22.20	9	8.30e-05	-2.9226	CDMS	
	N ₂ D ⁺	231320	3(3,3) - 2(3,3)	22.20	7	4.55e-05	-3.2931	CDMS	
	N ₂ D ⁺	231321	3(2,2) - 2(1,2)	22.20	5	1.39e-04	-2.9555	CDMS	
	N ₂ D ⁺	231321	3(2,2) - 2(1,1)	22.20	5	4.31e-04	-2.463	CDMS	
	N ₂ D ⁺	231321	3(2,1) - 2(1,0)	22.20	3	3.33e-04	-2.7964	CDMS	
	N ₂ D ⁺	231321	3(4,4) - 2(3,4)	22.20	9	3.59e-05	-3.2863	CDMS	
	N ₂ D ⁺	231322	3(3,3) - 2(2,3)	22.20	7	9.01e-05	-2.9965	CDMS	
	N ₂ D ⁺	231322	3(2,3) - 2(1,2)	22.20	7	6.07e-04	-2.1678	CDMS	
	N ₂ D ⁺	231322	3(3,3) - 2(2,2)	22.20	7	5.69e-04	-2.1961	CDMS	
	N ₂ D ⁺	231322	3 - 2	22.20	63	4.38e-04	-1.3554	CDMS	
	N ₂ D ⁺	231322	3(3,2) - 2(2,1)	22.20	5	5.49e-04	-2.3578	CDMS	
	N ₂ D ⁺	231322	3(3,4) - 2(2,3)	22.20	9	6.24e-04	-2.0467	CDMS	
	N ₂ D ⁺	231322	3(4,4) - 2(3,3)	22.20	9	6.78e-04	-2.0109	CDMS	
	N ₂ D ⁺	231322	3(4,3) - 2(3,2)	22.20	7	6.46e-04	-2.141	CDMS	
	N ₂ D ⁺	231322	3(4,5) - 2(3,4)	22.20	11	7.14e-04	-1.9011	CDMS	
	N ₂ D ⁺	231322	3(2,1) - 2(1,1)	22.20	3	2.81e-04	-2.8709	CDMS	
	N ₂ D ⁺	231322	3(3,2) - 2(2,2)	22.20	5	6.64e-05	-3.2752	CDMS	

Table C1 continued

Table C1 (continued)

Name	Formula	f_{rest} (MHz)	Transition Number	E_u (K)	g_u	A_u (s^{-1})	$I_{\text{CDMS,JPL}}$	Ref.
Ethanol	$N_2 D^+$	231322	3(4,3) - 2(3,3)	22.20	7	6.57e-05	-3.1332	CDMS
	C_2H_5OH	215890	25(4,22)t - 25(3,23)t	294.18	51	6.14e-05	...	JPL
	C_2H_5OH	215932	11(0,11)g- - 10(1,9)g+	115.30	23	3.09e-05	...	JPL
	C_2H_5OH	216127	12(9,3)t - 13(8,6)t	167.06	25	3.86e-06	...	JPL
	C_2H_5OH	216127	12(9,4)t - 13(8,5)t	167.06	25	3.86e-06	...	JPL
	C_2H_5OH	216273	40(5,35)t - 40(4,36)t	732.02	81	6.21e-05	...	JPL
	C_2H_5OH	216416	13(0,13)g+ - 12(0,12)g+	130.63	27	9.08e-05	...	JPL
	C_2H_5OH	216522	8(4,4)g+ - 7(3,4)g-	106.30	17	3.48e-05	...	JPL
	C_2H_5OH	216660	8(4,5)g+ - 7(3,5)g-	106.30	17	3.48e-05	...	JPL
	C_2H_5OH	216879	47(8,40)g+ - 47(7,40)g-	1075.45	95	3.51e-05	...	JPL
	C_2H_5OH	217100	45(6,39)t - 45(5,40)t	927.02	91	6.27e-05	...	JPL
	C_2H_5OH	217236	51(1,51)g+ - 50(2,49)g-	1098.44	103	1.69e-05	...	JPL
	C_2H_5OH	217238	51(0,51)g+ - 50(1,49)g-	1098.44	103	3.51e-05	...	JPL
	C_2H_5OH	217238	40(9,31)t - 39(10,30)t	791.86	81	1.83e-05	...	JPL
	C_2H_5OH	217400	31(6,26)t - 30(7,23)t	463.51	63	1.96e-05	...	JPL
	C_2H_5OH	217497	13(1,13)g- - 12(0,12)g-	134.51	27	3.36e-07	...	JPL
	C_2H_5OH	217497	34(5,29)g+ - 34(4,31)g-	586.69	69	3.57e-05	...	JPL
	C_2H_5OH	218461	5(3,2)t - 4(2,3)t	23.89	11	6.60e-05	...	JPL
	C_2H_5OH	218554	21(5,16)t - 21(4,17)t	226.01	43	6.21e-05	...	JPL
	C_2H_5OH	218729	56(13,44)g+ - 55(14,42)g-	1590.39	113	1.05e-05	...	JPL
	C_2H_5OH	218943	35(4,31)t - 35(3,32)t	559.69	71	6.47e-05	...	JPL
	C_2H_5OH	219067	22(4,19)g- - 22(3,19)g+	292.05	45	3.62e-05	...	JPL
	C_2H_5OH	219069	32(7,25)g+ - 32(6,27)g-	556.35	65	3.56e-05	...	JPL
	C_2H_5OH	219174	30(3,27)t - 30(2,28)t	409.96	61	6.51e-05	...	JPL
	C_2H_5OH	219241	52(13,40)t - 51(14,37)t	1369.56	105	1.78e-05	...	JPL
	C_2H_5OH	219241	52(13,39)t - 51(14,38)t	1369.56	105	1.78e-05	...	JPL
	C_2H_5OH	230145	25(3,23)t - 25(2,24)t	283.82	51	7.51e-05	...	JPL
	C_2H_5OH	230231	13(2,11)g- - 12(2,10)g-	143.28	27	1.07e-04	...	JPL
	C_2H_5OH	230673	13(2,11)g+ - 12(2,10)g+	138.62	27	1.07e-04	...	JPL
	C_2H_5OH	230794	6(5,1)g+ - 5(4,1)g-	104.80	13	6.12e-05	...	JPL
	C_2H_5OH	230794	6(5,2)g+ - 5(4,2)g-	104.80	13	6.12e-05	...	JPL
	C_2H_5OH	230954	16(5,11)t - 16(4,12)t	145.77	33	7.10e-05	...	JPL
	C_2H_5OH	230991	14(0,14)t - 13(1,13)t	85.53	29	6.87e-05	...	JPL
	C_2H_5OH	230997	41(5,36)t - 41(4,37)t	767.38	83	7.58e-05	...	JPL
	C_2H_5OH	231221	10(3,7)g- - 10(2,9)g+	118.42	21	4.12e-05	...	JPL
	C_2H_5OH	231277	55(14,41)t - 54(15,40)t	1538.49	111	2.07e-05	...	JPL
	C_2H_5OH	231277	55(14,42)t - 54(15,39)t	1538.49	111	2.07e-05	...	JPL
	C_2H_5OH	231559	21(5,17)t - 21(4,18)t	225.95	43	7.39e-05	...	JPL
	C_2H_5OH	231561	20(5,16)t - 20(4,17)t	208.21	41	7.36e-05	...	JPL
	C_2H_5OH	231669	14(1,14)g+ - 13(1,13)g+	141.90	29	1.11e-04	...	JPL
	C_2H_5OH	231738	19(5,15)t - 19(4,16)t	191.32	39	7.33e-05	...	JPL
	C_2H_5OH	231790	31(5,27)g- - 31(4,27)g+	505.98	63	4.30e-05	...	JPL
	C_2H_5OH	231790	22(5,18)t - 22(4,19)t	244.54	45	7.44e-05	...	JPL
	C_2H_5OH	231841	35(9,26)g- - 34(10,24)g+	684.80	71	1.15e-05	...	JPL
	C_2H_5OH	231874	20(10,10)g- - 21(9,12)g+	358.85	41	6.74e-06	...	JPL
	C_2H_5OH	231874	20(10,11)g- - 21(9,13)g+	358.85	41	6.74e-06	...	JPL
C_2H_5OH	232076	15(5,10)t - 15(4,11)t	132.29	31	7.13e-05	...	JPL	
C_2H_5OH	232287	29(4,26)g- - 29(3,27)g-	443.72	59	3.89e-07	...	JPL	
C_2H_5OH	232318	23(5,19)t - 23(4,20)t	263.99	47	7.52e-05	...	JPL	
C_2H_5OH	232367	33(6,27)g+ - 33(5,28)g+	565.06	67	3.84e-07	...	JPL	
C_2H_5OH	232405	17(5,13)t - 17(4,14)t	160.10	35	7.30e-05	...	JPL	
C_2H_5OH	232423	43(8,35)g+ - 42(9,33)g-	923.33	87	1.41e-05	...	JPL	
C_2H_5OH	232491	14(0,14)g+ - 13(0,13)g+	141.79	29	1.13e-04	...	JPL	
C_2H_5OH	232597	14(1,14)g- - 13(1,13)g-	146.53	29	1.12e-04	...	JPL	
C_2H_5OH	232809	16(5,12)t - 16(4,13)t	145.77	33	7.27e-05	...	JPL	
C_2H_5OH	232928	14(5,9)t - 14(4,10)t	119.66	29	7.11e-05	...	JPL	
C_2H_5OH	233209	24(5,20)t - 24(4,21)t	284.28	49	7.63e-05	...	JPL	
C_2H_5OH	233215	15(5,11)t - 15(4,12)t	132.29	31	7.23e-05	...	JPL	
C_2H_5OH	233229	43(20,24)g+ - 44(21,24)g-	1433.18	87	4.86e-05	...	JPL	
C_2H_5OH	233229	43(20,23)g+ - 44(21,23)g-	1433.18	87	4.86e-05	...	JPL	
C_2H_5OH	233571	13(5,8)t - 13(4,9)t	107.87	27	7.05e-05	...	JPL	
C_2H_5OH	233602	14(5,10)t - 14(4,11)t	119.65	29	7.17e-05	...	JPL	
Ethynyl (Deuterated)	CCD	216373	$N=3 - 2, J=7/2 - 5/2, F=9/2 - 7/2$	20.76	10	3.00e-05	-3.143	CDMS
	CCD	216373	$N=3 - 2, J=7/2 - 5/2, F=5/2 - 3/2$	20.76	6	2.67e-05	-3.4142	CDMS
	CCD	216373	$N=3 - 2, J=7/2 - 5/2, F=7/2 - 5/2$	20.76	8	2.76e-05	-3.275	CDMS
	CCD	216428	$N=3 - 2, J=5/2 - 3/2, F=7/2 - 5/2$	20.78	8	2.78e-05	-3.2724	CDMS
	CCD	216428	$N=3 - 2, J=5/2 - 3/2, F=5/2 - 3/2$	20.78	6	2.34e-05	-3.4728	CDMS
	CCD	216429	$N=3 - 2, J=5/2 - 3/2, F=3/2 - 1/2$	20.78	4	2.10e-05	-3.6956	CDMS
Formaldehyde	H_2CO	215976	11(2, 9) - 12(0,12)	279.73	23	3.15e-07	-5.5745	CDMS
	H_2CO	216569	9(1,8) - 9(1,9)	173.99	57	7.22e-06	-3.6678	CDMS
	H_2CO	218222	3(0,3) - 2(0,2)	20.95	7	2.82e-04	-2.769	CDMS
	H_2CO	218476	3(2,2) - 2(2,1)	68.10	7	1.57e-04	-3.0914	CDMS
	H_2CO	218760	3(2,1) - 2(2,0)	68.11	7	1.58e-04	-3.0903	CDMS

Table C1 continued

Table C1 (continued)

Name	Formula	f_{rest} (MHz)	Transition Number	E_u (K)	g_u	A_u (s^{-1})	$I_{\text{CDMS,JPL}}$	Ref.	
Formaldehyde (Doubly deuterated)	D ₂ CO	231410	4(0,4) - 3(0,3)	27.89	18	3.47e-04	-2.4242	CDMS	
	D ₂ CO	233650	4(2,3) - 3(2,2)	49.62	18	2.69e-04	-2.5711	CDMS	
Formamide	NH ₂ CHO	215688	10(2,8) - 9(2,7)	68.45	126	6.98e-04	-2.4894	JPL	
	NH ₂ CHO	216077	15(5,11) - 16(4,12)	196.78	93	5.62e-06	-4.6007	JPL	
	NH ₂ CHO	218460	10(1,9) - 9(1,8)	60.81	126	7.48e-04	-2.4541	JPL	
	NH ₂ CHO	232274	11(2,10) - 10(2,9)	78.95	138	8.82e-04	-2.3953	JPL	
	NH ₂ CHO	232798	20(6,14) - 21(5,17)	321.27	123	7.34e-06	-4.5759	JPL	
	NH ₂ CHO	233490	11(8, 3) - 10(8, 2), $F=11-10$	257.73	23	4.33e-04	-3.4424	JPL	
	NH ₂ CHO	233490	11(8, 4) - 10(8, 3), $F=11-10$	257.73	23	4.33e-04	-3.4424	JPL	
	NH ₂ CHO	233492	11(9, 2) - 10(9, 1), $F=10-9$	308.24	21	3.04e-04	-3.7092	JPL	
	NH ₂ CHO	233492	11(9, 3) - 10(9, 2), $F=10-9$	308.24	21	3.04e-04	-3.7092	JPL	
	NH ₂ CHO	233493	11(9, 2) - 10(9, 1), $F=11-10$	308.24	23	3.04e-04	-3.6693	JPL	
	NH ₂ CHO	233493	11(9, 3) - 10(9, 2), $F=11-10$	308.24	23	3.04e-04	-3.6693	JPL	
	NH ₂ CHO	233499	11(7, 4) - 10(7, 3), $F=11-10$	213.13	23	5.47e-04	-3.2763	JPL	
	NH ₂ CHO	233499	11(7, 5) - 10(7, 4), $F=11-10$	213.13	23	5.47e-04	-3.2763	JPL	
	NH ₂ CHO	233504	11(10,1) - 10(10,0), $F=10-9$	364.65	21	1.59e-04	-4.0706	JPL	
	NH ₂ CHO	233504	11(10,2) - 10(10,1), $F=10-9$	364.65	21	1.59e-04	-4.0706	JPL	
	NH ₂ CHO	233505	11(10,1) - 10(10,0), $F=11-10$	364.65	23	1.60e-04	-4.0308	JPL	
	NH ₂ CHO	233505	11(10,2) - 10(10,1), $F=11-10$	364.65	23	1.60e-04	-4.0308	JPL	
	NH ₂ CHO	233529	11(6,6) - 10(6,5)	174.45	92	6.51e-04	-2.6674	JPL	
	NH ₂ CHO	233529	11(6,5) - 10(6,4)	174.45	92	6.51e-04	-2.6674	JPL	
	NH ₂ CHO	233595	11(5,7) - 10(5,6)	141.71	92	7.36e-04	-2.5669	JPL	
NH ₂ CHO	233596	11(5,6) - 10(5,5)	141.71	92	7.36e-04	-2.5669	JPL		
NH ₂ CHO	233736	11(4,8) - 10(4,7)	114.93	69	8.07e-04	-2.4887	JPL		
NH ₂ CHO	233747	11(4,7) - 10(4,6)	114.93	69	8.07e-04	-2.4887	JPL		
NH ₂ CHO	233897	11(3,9) - 10(3,8)	94.11	138	8.62e-04	-2.43	JPL		
Formic Acid	HCOOH	218939	8(1,8) - 7(0,7)	40.52	17	2.46e-06	-5.1495	CDMS	
	HCOOH	231506	10(1,9) - 9(1,8)	64.47	21	1.33e-04	-3.1815	CDMS	
Formylium (Deuterated)	DCO ⁺	216113	19(2,18) - 18(2,17) E	20.74	7	7.66e-04	-3.9637	JPL	
Hydrogen Cyanide (Deuterated)	DCN	217237	$J=3-2, F=3-3$	20.86	7	5.08e-05	-2.7679	JPL	
	DCN	217238	$J=3-2, F=2-1$	20.86	5	3.84e-04	-2.0355	JPL	
	DCN	217239	$J=3-2, F=3-2$	20.86	7	4.07e-04	-1.8648	JPL	
	DCN	217239	$J=3-2, F=4-3$	20.86	21	4.58e-04	-1.7046	JPL	
	DCN	217241	$J=3-2, F=2-2$	20.86	5	7.12e-05	-2.7679	JPL	
Hydrogen Sulfide	H ₂ S	216710	2(2,0) - 2(1,1)	83.98	5	4.87e-05	-3.0171	CDMS	
	HNCO	218981	10(1,10) - 9(1,9)	101.08	21	1.42e-04	-2.6771	CDMS	
Isocyanic Acid	HNCO	219657	10(3,8) - 9(3,7)	432.96	21	1.20e-04	-2.6771	CDMS	
	HNCO	219657	10(3,7) - 9(3,6)	432.96	21	1.20e-04	-3.2322	CDMS	
	HNCO	219734	10(2,9) - 9(2,8)	228.29	21	1.35e-04	-2.8867	CDMS	
	HNCO	219737	10(2,8) - 9(2,7)	228.29	21	1.35e-04	-2.8867	CDMS	
	HNCO	219798	10(0,10) - 9(0,9)	58.02	21	1.47e-04	-2.6022	CDMS	
	HNCO	231873	28(1,28) - 29(0,29)	469.87	57	6.68e-05	-3.1298	CDMS	
	Methanol	CH ₃ OH	216946	5(1,4) - 4(2,3) E	55.87	44	1.21e-05	-4.5032	JPL
		CH ₃ OH	218440	4(2,2) - 3(1,2) E	45.46	36	4.69e-05	-3.9915	JPL
CH ₃ OH		230369	22(4,18) - 21(5,16) E	682.75	180	2.08e-05	-4.5906	JPL	
CH ₃ OH		231281	10(2,9) - 9(3,6) A2	165.35	84	1.83e-05	-4.2284	JPL	
CH ₃ OH		232419	10(2,8) - 9(3,7) A1	165.40	84	1.87e-05	-4.2242	JPL	
CH ₃ OH		232784	18(3,16) - 17(4,13) A1	446.53	148	2.16e-05	-4.3202	JPL	
CH ₃ OH		232945	10(3,7) - 11(2,9) E	190.37	84	2.13e-05	-4.2038	JPL	
CH ₃ OH		233796	18(3,15) - 17(4,14) A2	446.58	148	2.20e-05	-4.3164	JPL	
Methanol (Deuterated)		CH ₂ DOH	216130	12(0,12)e0 - 11(1,11)e1	166.52	25	1.55e-05	-4.269	JPL
		CH ₂ DOH	217072	3(2,2)e1 - 3(2,1)e0	39.43	7	4.00e-07	-6.2276	JPL
	CH ₂ DOH	217190	16(1,16)e1 - 15(3,13)e0	304.06	33	9.69e-08	-6.5535	JPL	
	CH ₂ DOH	217492	7(3,4)e1 - 8(1,7)e1	104.88	15	1.42e-08	-7.4418	JPL	
	CH ₂ DOH	218316	5(2,4)e1 - 5(1,5)e1	58.71	11	6.69e-06	-4.8387	JPL	
	CH ₂ DOH	219551	5(1,5)e1 - 4(1,4)e1	48.23	11	5.09e-06	-4.9441	JPL	
	CH ₂ DOH	230254	16(2,14)e0 - 16(1,15)e0	310.21	33	6.50e-05	-3.7605	JPL	
	CH ₂ DOH	230377	15(2,13)o1 - 15(1,14)o1	292.77	31	2.28e-05	-4.2186	JPL	
	CH ₂ DOH	231840	10(1,9)e0 - 9(2,8)e0	123.70	21	1.22e-05	-4.4162	JPL	
	CH ₂ DOH	232929	3(3,0)e0 - 4(2,3)e0	48.77	7	3.26e-06	-5.3608	JPL	
	CH ₂ DOH	233083	5(3,2)e0 - 4(2,2)o1	68.10	11	2.33e-06	-5.339	JPL	
	CH ₂ DOH	233134	5(3,3)e0 - 4(2,3)o1	68.09	11	2.35e-06	-5.3351	JPL	
	CH ₂ DOH	233142	5(3,3)e0 - 4(2,3)o1	260.58	29	2.22e-05	-5.3351	JPL	
	CH ₂ DOH	233210	17(2,15)e0 - 17(1,16)e0	347.06	35	6.49e-05	-3.7945	JPL	
	CH ₂ DOH	233461	9(2,8)e1 - 9(1,9)e1	122.94	19	8.68e-06	-4.6097	JPL	
	CH ₂ DOH	233736	7(3,5)e1 - 8(2,6)e0	104.88	15	2.70e-06	-5.1942	JPL	
	Methanol (¹³ C-substituted)	¹³ CH ₃ OH	215708	8(4,5) - 9(3, 6) A2	162.33	17	1.13e-05	-3.02967	CDMS
¹³ CH ₃ OH		215722	8(4,4) - 9(3, 7) A1	162.33	17	1.13e-05	-3.02974	CDMS	
¹³ CH ₃ OH		215887	4(2,2) - 3(1,2) E2	45.01	9	4.53e-05	-2.53306	CDMS	
¹³ CH ₃ OH		216370	10(2,9) - 9(3,6) A2	162.36	21	1.50e-05	-2.81704	CDMS	
¹³ CH ₃ OH		217045	14(1,13) - 13(2,12) A2	254.26	29	2.37e-05	-2.61267	CDMS	
¹³ CH ₃ OH		217400	10(2,8) - 9(3,7) A1	162.41	21	1.53e-05	-2.81298	CDMS	

Table C1 continued

Table C1 (*continued*)

Name	Formula	f_{rest} (MHz)	Transition Number	E_u (K)	g_u	A_u (s^{-1})	$I_{\text{CDMS,JPL}}$	Ref.
Silicon Monoxide	SiO	217105	5 - 4	31.26	11	5.21e-04	-1.3166	JPL
Sulfur Dioxide	SO ₂	216643	22(2,20) - 22(1,21)	248.44	45	9.27e-05	-3.0823	JPL
Sulfur Monoxide (³⁴ S-substituted)	³⁴ SO	215840	6(5) - 5(4)	34.38	13	1.26e-04	-2.3335	JPL

NOTE— The f_{rest} is the rest frequency in MHz, the E_u is the energy of the upper state in K, the g_u is the degeneracy factor, the A_u is the Einstein A coefficient in s^{-1} , the $I_{\text{CDMS,JPL}}$ is the CDMS/JPL intensity, the transition number is in the form of upper state - lower state.

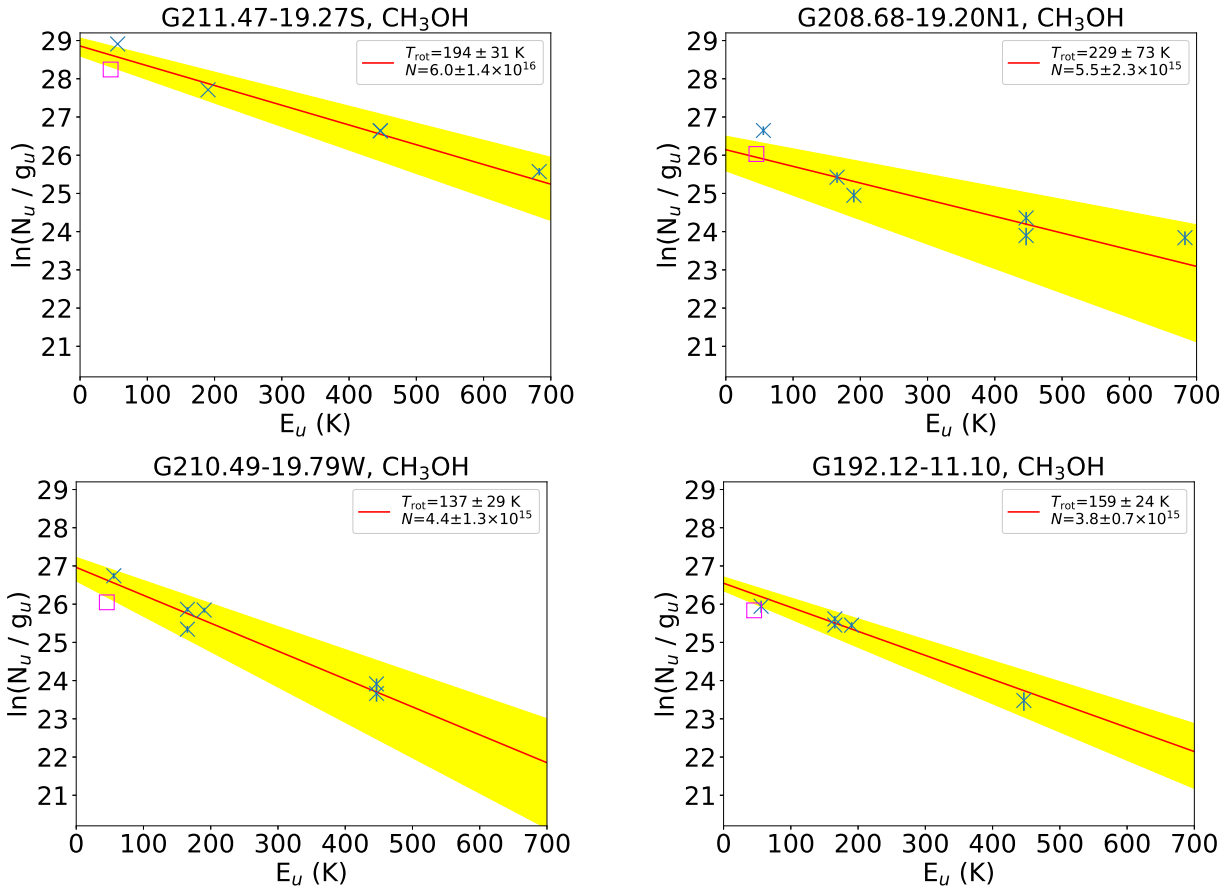
References— Jet Propulsion Laboratory Molecular Spectroscopy (JPL, [Pickett et al. 1998](#)); Cologne Database for Molecular Spectroscopy (CDMS, [Mller et al. 2005](#)); eXtended CASA Line Analysis Software Suite (XCLASS, [Möller et al. 2017](#))

D. POPULATION DIAGRAM OF METHANOL

The population diagram method is also a common tool for analyzing the excitation of molecular emission ([Goldsmith & Langer 1999](#)). In this section, we show the population diagrams of CH₃OH in all the four sources. We have excluded in the analysis the transition $J(K_a, K_c)=4(2,2)-3(1,2)$ E ($f_{\text{rest}}=218440$ MHz; $E_u=45$ K) since it has been reported to be possibly masing in different astrophysical objects ([Hunter et al. 2014](#); [Leurini et al. 2016](#); [Yuan et al. 2017](#)).

The total column density (N_{tot}) and the rotational temperature derived from the rotational diagrams are in general consistent with the MAGIX results (see Sect. 3.2). The data points in G208 at lower (< 200 K) and higher (> 200 K) energy levels seem to have different slopes. This could indicate the possibility of multiple molecular emission components in different excitation temperatures (e.g. cold envelope and hot corino).

Figure D1. The rotational diagrams (RD) of CH₃OH. The red line shows the best-fit result from the RD analysis and the yellow area represents the 1 σ interval. We adopted the molecular parameters, including the partition function, from Cologne Database for Molecular Spectroscopy (CDMS, Miller et al. 2005). The magenta squares mark the transition $J(K_a, K_c)=4(2,2)-3(1,2)$ E which is not used in the fitting (See Appendix D).



REFERENCES

- Araki, M., Takano, S., Sakai, N., et al. 2016, *ApJ*, 833, 291, doi: [10.3847/1538-4357/833/2/291](https://doi.org/10.3847/1538-4357/833/2/291)
- Araya, E., Hofner, P., Kurtz, S., Olmi, L., & Linz, H. 2008, *The Astrophysical Journal*, 675, 420, doi: [10.1086/527284](https://doi.org/10.1086/527284)
- Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, 558, A33, doi: [10.1051/0004-6361/201322068](https://doi.org/10.1051/0004-6361/201322068)
- Barone, V., Latouche, C., Skouteris, D., et al. 2015, *MNRAS*, 453, L31, doi: [10.1093/mnras/slv094](https://doi.org/10.1093/mnras/slv094)
- Beckwith, S. V. W., Sargent, A. I., Chini, R. S., & Guesten, R. 1990, *AJ*, 99, 924, doi: [10.1086/115385](https://doi.org/10.1086/115385)
- Belloche, A., Maury, A. J., Maret, S., et al. 2020, arXiv e-prints, arXiv:2002.00592. <https://arxiv.org/abs/2002.00592>
- Bergner, J. B., Martín-Doménech, R., Öberg, K. I., et al. 2019, *ACS Earth and Space Chemistry*, 3, 1564, doi: [10.1021/acsearthspacechem.9b00059](https://doi.org/10.1021/acsearthspacechem.9b00059)
- Bianchi, E., Codella, C., Ceccarelli, C., et al. 2017, *MNRAS*, 467, 3011, doi: [10.1093/mnras/stx252](https://doi.org/10.1093/mnras/stx252)
- . 2018, *MNRAS*, 483, 1850, doi: [10.1093/mnras/sty2915](https://doi.org/10.1093/mnras/sty2915)
- Bottinelli, S., Ceccarelli, C., Williams, J. P., & Lefloch, B. 2007, *A&A*, 463, 601, doi: [10.1051/0004-6361:20065139](https://doi.org/10.1051/0004-6361:20065139)
- Bottinelli, S., Ceccarelli, C., Lefloch, B., et al. 2004, *ApJ*, 615, 354, doi: [10.1086/423952](https://doi.org/10.1086/423952)
- Cazaux, S., Tielens, A. G. G. M., Ceccarelli, C., et al. 2003, *ApJ*, 593, L51, doi: [10.1086/378038](https://doi.org/10.1086/378038)
- Ceccarelli, C. 2004, in *Star Formation in the Interstellar Medium: In Honor of David Hollenbach, Chris McKee and Frank Shu*, ASP Conference Proceeding, ed. D. L. D. N. D. Johnstone, F.C. Adams & E. Ostriker, Vol. 323, San Francisco, Astronomical Society of the Pacific (ASP Conference Proceedings)
- Chini, R., Reipurth, B., Ward-Thompson, D., et al. 1997, *The Astrophysical Journal*, 474, L135, doi: [10.1086/310436](https://doi.org/10.1086/310436)
- Codella, C., et al. 2016, *A&A*, 586, L3, doi: [10.1051/0004-6361/201527424](https://doi.org/10.1051/0004-6361/201527424)
- Drozdovskaya, M. N., van Dishoeck, E. F., Jørgensen, J. K., et al. 2018, *MNRAS*, 476, 4949, doi: [10.1093/mnras/sty462](https://doi.org/10.1093/mnras/sty462)
- Eden, D. J., Liu, T., Kim, K.-T., et al. 2019, *MNRAS*, 485, 2895, doi: [10.1093/mnras/stz574](https://doi.org/10.1093/mnras/stz574)
- Esplugues, G. B., Viti, S., Goicoechea, J. R., & Cernicharo, J. 2014, *A&A*, 567, A95, doi: [10.1051/0004-6361/201323010](https://doi.org/10.1051/0004-6361/201323010)
- Evans, N. J., Francesco, J. D., Lee, J.-E., et al. 2015, *The Astrophysical Journal*, 814, 22, doi: [10.1088/0004-637x/814/1/22](https://doi.org/10.1088/0004-637x/814/1/22)
- Fang, M., Kim, J. S., van Boekel, R., et al. 2013, *ApJS*, 207, 5, doi: [10.1088/0067-0049/207/1/5](https://doi.org/10.1088/0067-0049/207/1/5)
- Fischer, W. J., Megeath, S. T., Ali, B., et al. 2010, *A&A*, 518, L122, doi: [10.1051/0004-6361/201014636](https://doi.org/10.1051/0004-6361/201014636)
- Friberg, P., Madden, S. C., Hjalmarsen, A., & Irvine, W. M. 1988, *A&A*, 195, 281
- Furlan, E., Fischer, W. J., Ali, B., et al. 2016, *ApJS*, 224, 5, doi: [10.3847/0067-0049/224/1/5](https://doi.org/10.3847/0067-0049/224/1/5)
- Garrod, R. T., Weaver, S. L. W., & Herbst, E. 2008, *ApJ*, 682, 283, doi: [10.1086/588035](https://doi.org/10.1086/588035)
- Goldsmith, P. F., & Langer, W. D. 1999, *ApJ*, 517, 209, doi: [10.1086/307195](https://doi.org/10.1086/307195)
- Halfen, D. T., Ilyushin, V., & Ziurys, L. M. 2011, *ApJ*, 743, 60, doi: [10.1088/0004-637X/743/1/60](https://doi.org/10.1088/0004-637X/743/1/60)
- Hatchell, J., Thompson, M. A., Millar, T. J., & MacDonald, G. H. 1998, *A&A*, 338, 713
- Herbst, E., & van Dishoeck, E. F. 2009, *ARA&A*, 47, 427, doi: [10.1146/annurev-astro-082708-101654](https://doi.org/10.1146/annurev-astro-082708-101654)
- Higuchi, A. E., Sakai, N., Watanabe, Y., et al. 2018, *ApJS*, 236, 52, doi: [10.3847/1538-4365/aabfe9](https://doi.org/10.3847/1538-4365/aabfe9)
- Hirano, N., & Liu, F.-C. 2014, *The Astrophysical Journal*, 789, 50, doi: [10.1088/0004-637x/789/1/50](https://doi.org/10.1088/0004-637x/789/1/50)
- Hunter, T. R., Brogan, C. L., Cyganowski, C. J., & Young, K. H. 2014, *ApJ*, 788, 187, doi: [10.1088/0004-637X/788/2/187](https://doi.org/10.1088/0004-637X/788/2/187)
- Imai, M., et al. 2016, *ApJ*, 830, L37, doi: [10.3847/2041-8205/830/2/L37](https://doi.org/10.3847/2041-8205/830/2/L37)
- Jacobsen, S. K., Jørgensen, J. K., Francesco, J. D., et al. 2019, *A&A*, 629, A29, doi: [10.1051/0004-6361/201833214](https://doi.org/10.1051/0004-6361/201833214)
- Jones, B. M., Bennett, C. J., & Kaiser, R. I. 2011, *ApJ*, 734, 78, doi: [10.1088/0004-637X/734/2/78](https://doi.org/10.1088/0004-637X/734/2/78)
- Jones, B. M., Bennett, C. J., & Kaiser, R. I. 2011, *ApJ*, 734, 78, doi: [10.1088/0004-637x/734/2/78](https://doi.org/10.1088/0004-637x/734/2/78)
- Jorgensen, J. K., Bourke, T. L., Myers, P. C., et al. 2005, *ApJ*, 632, 973, doi: [10.1086/433181](https://doi.org/10.1086/433181)
- Jørgensen, J. K., van der Wiel, M. H. D., Coutens, A., et al. 2016, *A&A*, 595, A117, doi: [10.1051/0004-6361/201628648](https://doi.org/10.1051/0004-6361/201628648)
- Jørgensen, J. K., Müller, H. S. P., Calcutt, H., et al. 2018, *A&A*, 620, A170, doi: [10.1051/0004-6361/201731667](https://doi.org/10.1051/0004-6361/201731667)
- Juvela, M., He, Jinhua, Pattle, Katherine, et al. 2018, *A&A*, 612, A71, doi: [10.1051/0004-6361/201731921](https://doi.org/10.1051/0004-6361/201731921)
- Kauffmann, J., Bertoldi, F., Bourke, T. L., Evans, N. J., & Lee, C. W. 2008, *A&A*, 487, 993, doi: [10.1051/0004-6361:200809481](https://doi.org/10.1051/0004-6361:200809481)
- Le Gal, R., berg, K. I., Loomis, R. A., Pegues, J., & Bergner, J. B. 2019, *The Astrophysical Journal*, 876, 72, doi: [10.3847/1538-4357/ab1416](https://doi.org/10.3847/1538-4357/ab1416)
- Lee, C.-F., Codella, C., Li, Z.-Y., & Liu, S.-Y. 2019, *ApJ*, 876, 63, doi: [10.3847/1538-4357/ab15db](https://doi.org/10.3847/1538-4357/ab15db)
- Lee, C.-F., Li, Z.-Y., Ho, P. T. P., et al. 2017, *ApJ*, 843, 27, doi: [10.3847/1538-4357/aa7757](https://doi.org/10.3847/1538-4357/aa7757)
- Lefloch, B., Bachiller, R., Ceccarelli, C., et al. 2018, *MNRAS*, 477, 4792, doi: [10.1093/mnras/sty937](https://doi.org/10.1093/mnras/sty937)
- Leurini, S., Menten, K. M., & Walmsley, C. M. 2016, *A&A*, 592, A31, doi: [10.1051/0004-6361/201527974](https://doi.org/10.1051/0004-6361/201527974)
- Liu, T., Kim, K.-T., Juvela, M., et al. 2018, *ApJS*, 234, 28, doi: [10.3847/1538-4365/aaa3dd](https://doi.org/10.3847/1538-4365/aaa3dd)

- López-Sepulcre, A., Sakai, N., Neri, R., et al. 2017, *A&A*, 606, A121, doi: [10.1051/0004-6361/201630334](https://doi.org/10.1051/0004-6361/201630334)
- Luo, G., Feng, S., Li, D., et al. 2019, *ApJ*, 885, 82, doi: [10.3847/1538-4357/ab45ef](https://doi.org/10.3847/1538-4357/ab45ef)
- Lpez-Sepulcre, A., Jaber, A. A., Mendoza, E., et al. 2015, *MNRAS*, 449, 2438, doi: [10.1093/mnras/stv377](https://doi.org/10.1093/mnras/stv377)
- Manoj, P., Watson, D. M., Neufeld, D. A., et al. 2013, *ApJ*, 763, 83, doi: [10.1088/0004-637x/763/2/83](https://doi.org/10.1088/0004-637x/763/2/83)
- Manoj, P., Green, J. D., Megeath, S. T., et al. 2016, *ApJ*, 831, 69, doi: [10.3847/0004-637X/831/1/69](https://doi.org/10.3847/0004-637X/831/1/69)
- Martín-Doménech, R., Bergner, J. B., berg, K. I., & Jørgensen, J. K. 2019, *ApJ*, 880, 130, doi: [10.3847/1538-4357/ab2a08](https://doi.org/10.3847/1538-4357/ab2a08)
- Martin-Domenech, R., Oberg, K. I., & Rajappan, M. 2020, Formation of NH₂CHO and CH₃CHO upon UV photoprocessing of interstellar ice analogs. <https://arxiv.org/abs/2004.01151>
- McMullin, J. P., W. B. S. D. Y. W. . G. K. 2007, in *Astronomical Data Analysis Software and Systems XVI*, ed. F. H. R. A. Shaw & D. J. Bell, Vol. 376, San Francisco, Astronomical Society of the Pacific (ASP Conference Series)
- Megeath, S. T., Gutermuth, R., Muzerolle, J., et al. 2012, *AJ*, 144, 192, doi: [10.1088/0004-6256/144/6/192](https://doi.org/10.1088/0004-6256/144/6/192)
- Möller, T., Endres, C., & Schilke, P. 2017, *A&A*, 598, A7, doi: [10.1051/0004-6361/201527203](https://doi.org/10.1051/0004-6361/201527203)
- Möller, T., Bernst, I., Panoglou, D., et al. 2013, *A&A*, 549, A21, doi: [10.1051/0004-6361/201220063](https://doi.org/10.1051/0004-6361/201220063)
- Miller, H. S., Schlder, F., Stutzki, J., & Winnenwiser, G. 2005, *Journal of Molecular Structure*, 742, 215, doi: <https://doi.org/10.1016/j.molstruc.2005.01.027>
- Noble, J. A., Theule, P., Congiu, E., et al. 2015, *A&A*, 576, A91, doi: [10.1051/0004-6361/201425403](https://doi.org/10.1051/0004-6361/201425403)
- Ospina-Zamudio, J., Lefloch, B., Ceccarelli, C., et al. 2018, *A&A*, 618, A145, doi: [10.1051/0004-6361/201832857](https://doi.org/10.1051/0004-6361/201832857)
- Oya, Y., Sakai, N., Watanabe, Y., et al. 2017, *ApJ*, 837, 174, doi: [10.3847/1538-4357/aa6300](https://doi.org/10.3847/1538-4357/aa6300)
- Palumbo, M. E., Geballe, T. R., & Tielens, A. G. G. M. 1997, *ApJ*, 479, 839, doi: [10.1086/303905](https://doi.org/10.1086/303905)
- Parise, B., Ceccarelli, C., Tielens, A. G. G. M., et al. 2006, *A&A*, 453, 949, doi: [10.1051/0004-6361:20054476](https://doi.org/10.1051/0004-6361:20054476)
- Persson, M. V., Jørgensen, J. K., Müller, H. S. P., et al. 2018, *A&A*, 610, A54, doi: [10.1051/0004-6361/201731684](https://doi.org/10.1051/0004-6361/201731684)
- Pickett, H., Poynter, R., Cohen, E., et al. 1998, *Journal of Quantitative Spectroscopy and Radiative Transfer*, 60, 883, doi: [https://doi.org/10.1016/S0022-4073\(98\)00091-0](https://doi.org/10.1016/S0022-4073(98)00091-0)
- Planck Collaboration, et al. 2016, *A&A*, 594, A28, doi: [10.1051/0004-6361/201525819](https://doi.org/10.1051/0004-6361/201525819)
- Price-Whelan, A. M., Sipőcz, B. M., Günther, H. M., et al. 2018, *AJ*, 156, 123, doi: [10.3847/1538-3881/aabc4f](https://doi.org/10.3847/1538-3881/aabc4f)
- Quan, D., & Herbst, E. 2007, *A&A*, 474, 521, doi: [10.1051/0004-6361:20078246](https://doi.org/10.1051/0004-6361:20078246)
- Raunier, S., Chiavassa, T., Duvernay, F., et al. 2004, *A&A*, 416, 165, doi: [10.1051/0004-6361:20034558](https://doi.org/10.1051/0004-6361:20034558)
- Roberts, H., & Millar, T. J. 2000, *A&A*, 361, 388
- Sahu, D., Liu, S.-Y., Su, Y.-N., et al. 2019, *The Astrophysical Journal*, 872, 196, doi: [10.3847/1538-4357/aaffda](https://doi.org/10.3847/1538-4357/aaffda)
- Sakai, N., Sakai, T., Hirota, T., Burton, M., & Yamamoto, S. 2009, *ApJ*, 697, 769, doi: [10.1088/0004-637x/697/1/769](https://doi.org/10.1088/0004-637x/697/1/769)
- Sakai, N., Sakai, T., Hirota, T., & Yamamoto, S. 2008, *ApJ*, 672, 371, doi: [10.1086/523635](https://doi.org/10.1086/523635)
- Saladino, R., Botta, G., Pino, S., Costanzo, G., & Di Mauro, E. 2012, *Chem. Soc. Rev.*, 41, 5526, doi: [10.1039/C2CS35066A](https://doi.org/10.1039/C2CS35066A)
- Sanhueza, P., Jackson, J. M., Foster, J. B., et al. 2013, *ApJ*, 773, 123, doi: [10.1088/0004-637X/773/2/123](https://doi.org/10.1088/0004-637X/773/2/123)
- Takahashi, S., & Ho, P. T. P. 2012, *The Astrophysical Journal*, 745, L10, doi: [10.1088/2041-8205/745/1/L10](https://doi.org/10.1088/2041-8205/745/1/L10)
- Takahashi, S., Machida, M. N., Tomisaka, K., et al. 2019, *The Astrophysical Journal*, 872, 70, doi: [10.3847/1538-4357/aaf6ed](https://doi.org/10.3847/1538-4357/aaf6ed)
- Taquet, V., Charnley, S. B., & Sipil, O. 2014, *ApJ*, 791, 1, doi: [10.1088/0004-637x/791/1/1](https://doi.org/10.1088/0004-637x/791/1/1)
- Taquet, V., López-Sepulcre, A., Ceccarelli, C., et al. 2015, *ApJ*, 804, 81, doi: [10.1088/0004-637x/804/2/81](https://doi.org/10.1088/0004-637x/804/2/81)
- Tatematsu, K., Liu, T., Ohashi, S., et al. 2017, *ApJS*, 228, 12, doi: [10.3847/1538-4365/228/2/12](https://doi.org/10.3847/1538-4365/228/2/12)
- Tobin, J. J., Megeath, S. T., van't Hoff, M., et al. 2019, *ApJ*, 886, 6, doi: [10.3847/1538-4357/ab498f](https://doi.org/10.3847/1538-4357/ab498f)
- van Kempen, T. A., Hogerheijde, M. R., van Dishoeck, E. F., et al. 2016, *A&A*, 587, A17, doi: [10.1051/0004-6361/201424725](https://doi.org/10.1051/0004-6361/201424725)
- Velusamy, T., Langer, W. D., & Thompson, T. 2014, *The Astrophysical Journal*, 783, 6, doi: [10.1088/0004-637x/783/1/6](https://doi.org/10.1088/0004-637x/783/1/6)
- Watson, D. M., Calvet, N. P., Fischer, W. J., et al. 2016, *ApJ*, 828, 52, doi: [10.3847/0004-637x/828/1/52](https://doi.org/10.3847/0004-637x/828/1/52)
- Wenger, M., Ochsenbein, F., Egret, D., et al. 2000, *A&AS*, 143, 9, doi: [10.1051/aas:2000332](https://doi.org/10.1051/aas:2000332)
- Wirström, E. S., Geppert, W. D., Hjalmarson, Å., et al. 2011, *A&A*, 533, A24, doi: [10.1051/0004-6361/201116525](https://doi.org/10.1051/0004-6361/201116525)
- Yang, Y.-L., Evans, Neal J., I., Smith, A., et al. 2020, *ApJ*, 891, 61, doi: [10.3847/1538-4357/ab7201](https://doi.org/10.3847/1538-4357/ab7201)
- Yi, H.-W., Lee, J.-E., Liu, T., et al. 2018, *ApJS*, 236, 51, doi: [10.3847/1538-4365/aac2e0](https://doi.org/10.3847/1538-4365/aac2e0)
- Yuan, J., Li, J.-Z., Wu, Y., et al. 2017, 852, 12, doi: [10.3847/1538-4357/aa9d40](https://doi.org/10.3847/1538-4357/aa9d40)