

## DIRECT SPECTRUM OF THE BENCHMARK T DWARF HD 19467 B

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## ABSTRACT

HD 19467 B is presently the only directly imaged T dwarf companion known to induce a measurable Doppler acceleration around a solar-type star. We present spectroscopy measurements of this important benchmark object taken with the Project 1640 integral field unit at Palomar Observatory. Our high-contrast  $R \approx 30$  observations obtained simultaneously across the  $JH$  bands confirm the cold nature of the companion as reported from the discovery article and determine its spectral type for the first time. Fitting the measured spectral energy distribution to SpeX/IRTF T dwarf standards and synthetic spectra from BT-Settl atmospheric models, we find that HD 19467 B is a  $T5.5 \pm 1$  dwarf with effective temperature  $T_{\text{eff}} = 978^{+20}_{-43}$  K. Our observations reveal significant methane absorption affirming its substellar nature. HD 19467 B shows promise to become the first T dwarf that simultaneously reveals its mass, age, and metallicity independent from the spectrum of light that it emits.

*Key words:* brown dwarfs – stars: individual (HD 19467) – techniques: high angular resolution – techniques: spectroscopic

## 1. INTRODUCTION

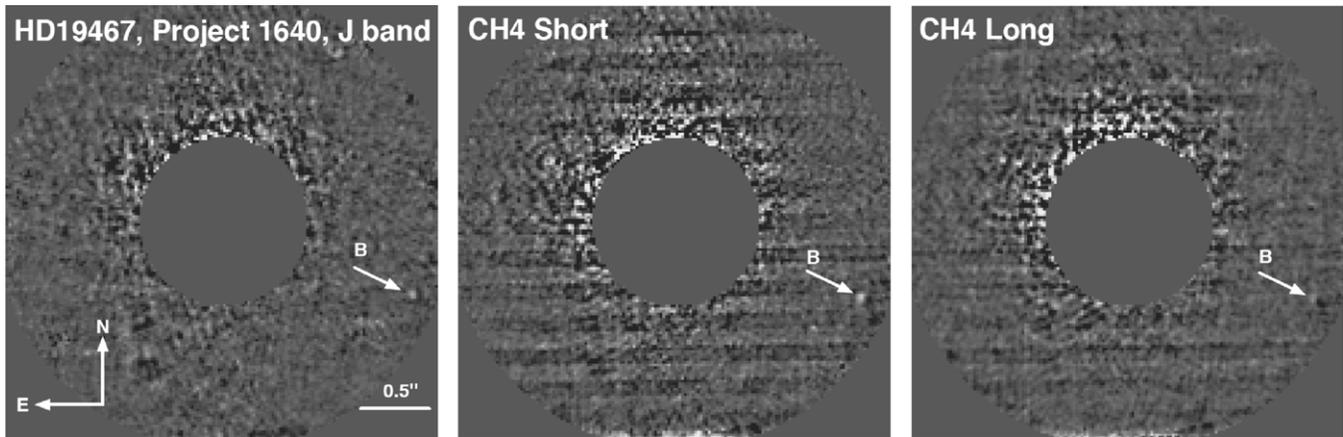
HD 19467 B is a faint co-moving companion to the nearby G3V star HD 19467 that was recently discovered as part of the TrenDS high-contrast imaging program (Crepp et al. 2012b, 2014). Prior to its direct imaging detection, the existence of HD 19467 B was first inferred from 16.9 years of precise stellar radial velocity (RV) measurements that revealed a long-term systemic acceleration of  $dv/dt = -1.37 \pm 0.09 \text{ m s}^{-1} \text{ yr}^{-1}$ . HD 19467 B's intrinsic brightness ( $M_J = 17.61 \pm 0.11$ ) and blue near-infrared colors ( $J - H = -0.29 \pm 0.15$ ,  $J - K_s = -0.36 \pm 0.14$ ) suggest that it is non-hydrogen-fusing. A lower mass limit of  $M \geq 51.9^{+3.6}_{-4.3} M_{\text{Jup}}$  derived from the RV acceleration and projected orbital separation of  $51.1 \pm 1.0 \text{ AU}$  is consistent with the model-dependent mass estimate,  $M = 56.7^{+4.6}_{-7.2} M_{\text{Jup}}$ , and supports the interpretation that HD 19467 B is a cold brown dwarf. Crepp et al. (2014) assigned a preliminary spectral-type of  $\approx T5\text{--}T7$  based on the object's location in color-magnitude diagrams but spectroscopic observations have yet to be reported.

HD 19467 B is an important benchmark object. It is currently the only directly imaged T dwarf companion known to cause a measurable Doppler acceleration around a nearby solar-type star. In addition to a precise parallax ( $32.40 \pm 0.62 \text{ mas}$ ) and many years of legacy RV measurements, astrometric observations already show systemic orbital motion at  $22 \pm 6 \text{ mas yr}^{-1}$  (van Leeuwen 2007). A tight mass constraint from dynamics will be available once both the astrometry and RV observations

reveal curvature (Crepp et al. 2012a). Further, the chemical composition of HD 19467 B has been inferred from its Sun-like (G3V) parent star. Crepp et al. (2014) find that HD 19467 A has a sub-solar metallicity of  $[\text{Fe}/\text{H}] = -0.15 \pm 0.04$ , thus HD 19467 B should also have relatively low metal content assuming it formed from the same cloud of material. Such model independent mass and metallicity measurements are essential for calibrating theoretical spectral and evolutionary models, particularly at cold temperatures.

The age of the HD 19467 system is currently less certain but forthcoming observations will place strong constraints on the primary star's evolutionary state. At a distance of only  $30.9 \pm 0.6 \text{ pc}$ , the interferometric CHARA array may be able to spatially resolve the surface of the star at near-infrared wavelengths, providing a direct measure of its radius (Boyajian et al. 2009). It is expected that the age could be tightly constrained because HD 19467 resides  $\Delta M_V = 0.28$  above the *Hipparcos* median main-sequence, has a subsolar metallicity, yet is nearly identical to the Sun. Therefore, we already know that HD 19467 is older than the Sun ( $>4.6 \text{ Gyr}$ ), and a radius constraint, even if only an upper limit (i.e., marginally spatially resolved by CHARA interferometry), will eliminate ages older than some threshold value.

HD 19467 B has an ideal separation ( $1''.65$ ) for follow-up integral field spectroscopy measurements (e.g., Bowler et al. 2010; Konopacky et al. 2013; Oppenheimer et al. 2013). In this Letter, we use the Project 1640 high-contrast instrument at Palomar to measure HD 19467 B's spectral-energy distribution



**Figure 1.** Reconnaissance images of HD 19467 B taken with the Project 1640 integral field unit at Palomar Observatory using the 200-inch telescope on 2013 October 17. High-contrast measurements in multiple bands recover the companion following PSF subtraction. A marked decline in the flux of HD 19467 B in the  $CH_4$  band ( $\lambda_{\text{short}} = 1.560\text{--}1.600 \mu\text{m}$ ,  $\lambda_{\text{long}} = 1.635\text{--}1.675 \mu\text{m}$ ) reveals significant methane absorption (see text for discussion). A description of  $H$ -band methane filter observations can be found in Janson et al. (2013).

in the  $JH$  bands and assess its physical properties. Spectral standard comparisons and theoretical model fits to the data allow us to assign a spectral type and directly study the object’s atmosphere for the first time.

## 2. PALOMAR HIGH-CONTRAST OBSERVATIONS

HD 19467 was observed on 2013 October 17 UT at the Palomar Observatory using the Hale 200 inch telescope. The PALM-3000 AO system provides diffraction-limited images at near-infrared wavelengths using a Shack-Hartmann wavefront sensor (Dekany et al. 2013). Given the brightness of HD 19467 ( $V = 7.0$ ), we used  $64 \times 64$  subapertures to provide fine spatial wavefront correction while running the AO system at 1 kHz. Simultaneous imaging and spectroscopy measurements were acquired by the Project 1640 (hereafter, P1640) high-contrast integral field spectrometer (IFS; Hinkley et al. 2011).

The seeing was estimated to be  $1''.3$  during the time of observations. We acquired 25 individual exposures totaling 4583 s of on-source integration time with the star centered behind the coronagraphic mask. Unocculted images were acquired to measure the flux ratio between the primary and secondary as a function of wavelength. HD 19467 B resides just inside of the P1640 field of view ( $4''.2 \times 4''.2$ ). The companion can be noticed when playing a “color movie” by sequentially viewing each image of the IFS data cube. Pre-speckle-suppressed images also show qualitatively that the companion appears to exhibit methane absorption in the  $H$ -band. We demonstrate that HD 19467 B must be a very cold object in the following analysis.

### 2.1. Spectral Extraction

Raw frames were converted into data cubes using the algorithm described in detail in Zimmerman et al. (2010). Each exposure creates 32 separate images taken at slightly different wavelengths (Hinkley et al. 2011). Although the companion light is initially mixed with stellar speckles, flux in each wavelength channel can be retrieved by performing point-spread function (PSF) subtraction that takes advantage of color information provided by the IFS (Crepp et al. 2011; Pueyo et al. 2012).

Once raw data frames were processed, we employed the techniques presented in Fergus et al. (2014) using the Spatial-Spectral model for Speckle Suppression (S4) algorithm. S4

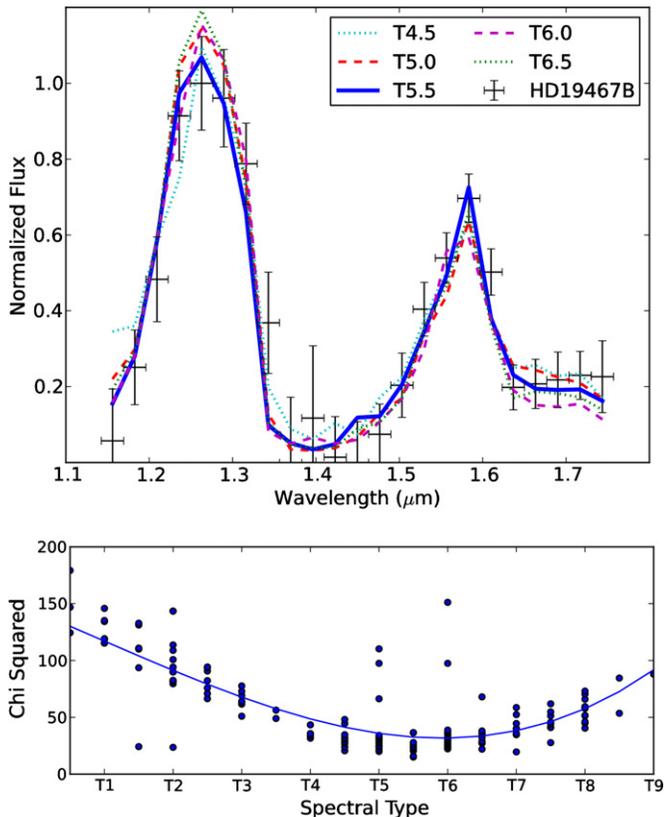
uses principal components analysis (PCA) to identify a linear combination of orthogonal (Karhunen-Loeve) modes to fit and remove the PSF of HD 19467 A in each wavelength channel (Figure 1). A concise explanation for how S4 is adapted to P1640 can be found in Appendix B of Oppenheimer et al. (2013). Various derivatives of PCA-based algorithms have been implemented previously for high-contrast imaging data (Amara & Quanz 2012). We find that S4 provides consistent and stable behavior for extracting the spectrum of synthetic and real companions irrespective of relative brightness and angular separation for P1640 data cubes (A. Veicht et al. 2015, in preparation).

HD 19467 B has an angular separation ( $1''.65$ ) and  $H$ -band contrast ratio ( $\Delta H = 12.46 \pm 0.10$ ) comparable to the planet HR 8799 b (Marois et al. 2008). Benefitting from correction of non-common-path errors from the P1640 internal wavefront calibration unit, our observations resulted in the companion being comparable in brightness to local speckles in frames spanning the  $JH$ -bands (Zhai et al. 2012; Cady et al. 2013). The relative flux of HD 19467 B was measured using aperture photometry with the occulted data cubes following speckle reduction by S4. Residual speckle noise was estimated by sampling many apertures at the same radial distance from the host star but at different azimuthal angles.

The extracted companion flux was calibrated using several unocculted short exposures of the host star before the observing sequence to derive a spectral response function (SRF). The SRF was used to scale the final spectrum accordingly, thus accounting for wavelength-dependent sky and instrument transmission. We validated the SRF using a nearby G3V star (HD 10697) from the IRTF spectral library. Additional calibration sources (HIP 112821, HIP 43567) were also observed before and after HD 19467 to track airmass dependent and temporal variations in the SRF. The high-resolution IRTF spectra were binned to match the wavelength sampling and coverage of the IFS. Comparing the SRF from HD 19467 to the above standard stars, we find that variations in the SRF are smaller than uncertainty in the companion flux introduced by the speckle suppression process. The fully calibrated spectrum of HD 19467 B is shown in Figure 2.

### 2.2. Astrometry

The same speckle suppression procedures may be employed to self-consistently extract a precise location for the companion



**Figure 2.** (Top) P1640 spectrum of HD 19467 B (black crosses) plotted with binned and trimmed SpeX/IRTF spectra of T4.5–T6.5 objects. (Bottom)  $\chi^2$  as a function of spectral type for T0–T9 objects and second-order polynomial fit. We derive a spectral type of  $T5.5 \pm 1$  from this spectral comparison.

measured relative to the star. From S4 processed images we find that HD 19467 B has an angular separation of  $\rho = 1'.640 \pm 0'.007$  and position angle  $PA = 241^\circ.7 \pm 0^\circ.3$ . These values are consistent with that from Crepp et al. (2014) and confirm the measured systemic (clockwise) astrometric orbital motion. A longer time baseline however is required to constrain the companion mass and orbital parameters at levels much beyond that reported in the discovery paper.

### 3. SPECTRAL TYPING

Crepp et al. (2014) estimate a spectral type of T5–T7 for HD 19467 B based on near-infrared broadband colors and  $JHK$  absolute magnitudes (Leggett et al. 2010; Dupuy and Liu 2012). We refine this analysis using the  $R = \lambda/\Delta\lambda \approx 30$  P1640 spectrum of HD 19467 B (Figure 2). T dwarfs are optimally assigned a spectral type by comparing their near-infrared (0.8–2.5  $\mu\text{m}$ ) spectra to spectral standards (Burgasser et al. 2006a). We compare the P1640 spectrum of HD 19467 B to 107 T dwarfs with spectral types ranging from T0 to T9. Spectra are from Burgasser et al. (2004, 2006a, 2008, 2010), Burgasser (2007), Chiu et al. (2006), Cruz et al. (2004), Kirkpatrick et al. (2011), Liebert & Burgasser (2007), Looper et al. (2007), Mace et al. (2013), Mainzer et al. (2011), and Sheppard & Cushing (2009). Most of the spectra were obtained via the SpeX prism library.<sup>12</sup>

We first bin and trim the SpeX/IRTF spectra to match the wavelength range and resolution of the observations. We then

calculate  $\chi^2$  for each binned T dwarf spectrum compared to the measurements;  $\chi^2$  is found by incorporating errors from both the P1640 spectrum of HD 19467 B and SpeX templates. A spectral type of T5.5 (2MASS J11101001+0116130, Burgasser et al. 2006c) produces the minimum  $\chi^2$  value. Figure 2 shows  $\chi^2$  as a function of spectral type for T0–T9 dwarfs with a second-order polynomial fit. Spectral types earlier than T4 all result in higher  $\chi^2$  values. We find a minimum (reduced) goodness-of-fit value of  $\chi_r^2 = 0.68$  with 22 degrees of freedom. From this analysis, we derive a spectral type of  $T5.5 \pm 1$ , compatible with the estimated T5–T7 spectral type from Crepp et al. (2014) which was based on broadband photometry alone.

Methane absorption represents another characteristic feature of cold brown dwarf atmospheres (Oppenheimer et al. 1995). In fact, L and T dwarf spectral sequences may be distinguished based on the existence of  $CH_4$  in the  $H$ -band. We estimate the  $CH_4$  spectral index of HD 19467 B using the following metric from Geballe et al. (2002):

$$S_{CH_4} = \frac{\int_{1.560}^{1.600} f_\lambda d\lambda}{\int_{1.635}^{1.675} f_\lambda d\lambda}, \quad (1)$$

where the integrals are carried out over the continuum (numerator) compared to wavelengths where  $CH_4$  is present in late-type T dwarfs (denominator). Brown dwarfs have  $S_{CH_4} > 1$ . The methane index increases monotonically for progressively lower temperature objects. Using  $f_\lambda$  values from Figure 2 we find that HD 19467 B has  $S_{CH_4} = 3.1^{+0.6}_{-0.8}$ . These values correspond to an empirically determined spectral type of  $\approx T6$  (Geballe et al. 2002). While the Geballe metric is meant for interpretation of higher resolution spectra than that provided by P1640, we use it here as a qualitative indicator to demonstrate that HD 19467 B shows significant methane absorption and the spectral type found from this specific index is compatible with fits to the broader spectrum.

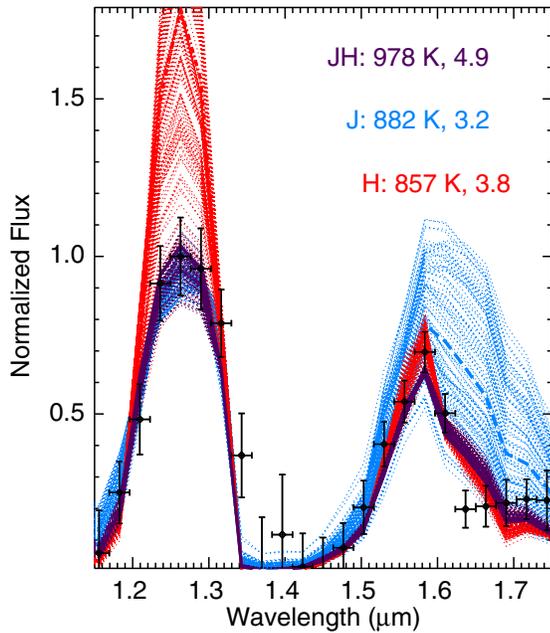
### 4. MODEL ATMOSPHERES

T dwarf spectral types are defined by near-infrared spectral morphology and do not necessarily correspond monotonically to physical properties (Kirkpatrick et al. 2008). However, by comparing the observed spectrum to synthetic spectra from model atmospheres, we can infer parameters such as effective temperature and surface gravity. We compare the P1640 spectrum of HD 19467 B to synthetic spectra from the most recent BT-Settl model atmospheres (Allard 2014). The model grid was evaluated at increments of 50–100 K from 400 to 4500 K for surface gravities  $\log(g) = 3.5, 4.0, 4.5, 5.0,$  and  $5.5$  (cgs units). All models are solar metallicity. Our fitting procedure is summarized below.

Model spectra are binned from their native resolution of  $\Delta\lambda = 1 \text{ \AA}$  to match that of P1640 ( $\Delta\lambda = 27 \text{ nm}$ ). A goodness-of-fit parameter ( $\chi^2$ ) is calculated for every binned spectrum in the model grid. We then generate probability distributions,  $P \propto \exp -\chi^2/2$ , using a  $10^6$  step Markov Chain Monte Carlo (MCMC) simulation with the Metropolis Hastings algorithm. The MCMC routine begins at the model with the smallest  $\chi^2$  value and interpolates between model results. We find that jump sizes of 100–250 K in temperature and 1.0–2.5 dex in surface gravity produce optimal acceptance ratios of  $\sim 0.35$ .

We fit four versions of the extracted P1640 spectrum of HD 19467 B: the complete  $JH$  spectrum, the  $JH$  spectrum without the four flux points closest to the water absorption band

<sup>12</sup> <http://pono.ucsd.edu/~adam/browndwarfs/speprism>

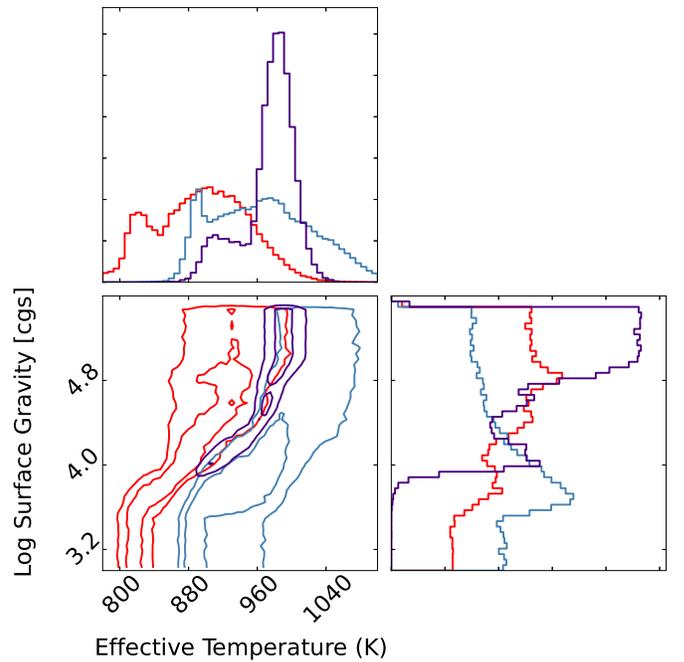


**Figure 3.** Best fit spectra from MCMC calculations for the complete *JH* (purple), *J*-band (blue) and *H*-band (red) P1640 spectrum of HD 19467 B (black points and error bars). Thin dotted lines are 100 spectra randomly selected from the posterior distributions plotted in Figure 4 to represent the range of model fits allowed within  $1\sigma$  uncertainties. Values listed in the legend indicate the mode of the temperature and surface gravity from the posterior distribution. HD 19467 B has a best fit effective temperature of  $T_{\text{eff}} = 978^{+20}_{-43}$  K.

at  $1.4\ \mu\text{m}$  (“trimmed” *JH*), and individual *J*-band and *H*-band spectra. Results for the complete versus trimmed *JH* spectra are effectively identical so we present results for the complete *JH* spectrum and the individual *J*- and *H*-band spectra (Figure 3). We note that P1640 acquires *JH* measurements simultaneously whereas most other high-contrast instruments must use separate *J*, *H* filters for non-contemporaneous observations. We have decided to compare derived effective temperatures for individual filters to highlight the fact that results can vary appreciably depending on bandpass.

Goodness-of-fit calculations ( $\chi^2$ ) resulted in a global minimum between  $700\text{K} \leq T_{\text{eff}} \leq 1100\text{K}$  at all surface gravities and using all four spectral fits. Posterior distributions in Figure 4 show multiple peaks in effective temperature with both the *J*-band (902–1033) and *H*-band (832–949 K) distributions producing slightly cooler results than the complete *JH* data set (934–998 K). Using the mode posterior value and 68% confidence interval from the *JH* spectrum, we find that HD 19467 B has an effective temperature of  $T_{\text{eff}} = 978^{+20}_{-43}$  K. This value is consistent with results for objects with similar spectral types from previous studies using bolometric luminosities (Golimowski et al. 2004; Vrba et al. 2004) and higher spectral resolution model comparisons (Burgasser et al. 2006b; del Burgo et al. 2009).

Posterior distributions for the surface gravity marginalized over effective temperature likewise show multiple peaks. The individual *J*-band and *H*-band results span two orders of magnitude but the combined *JH* results show convergence toward higher surface gravity,  $\log g = 4.21\text{--}5.31$ . To test the veracity of the surface gravity results, we fit model spectra to 15 T dwarfs from the SpeX Prism library binned and trimmed to match P1640. The template objects include nine spectral standards (T0 to T8), three additional late-type objects (T7.5, T8, and T8pec) and two early T dwarfs, including the young



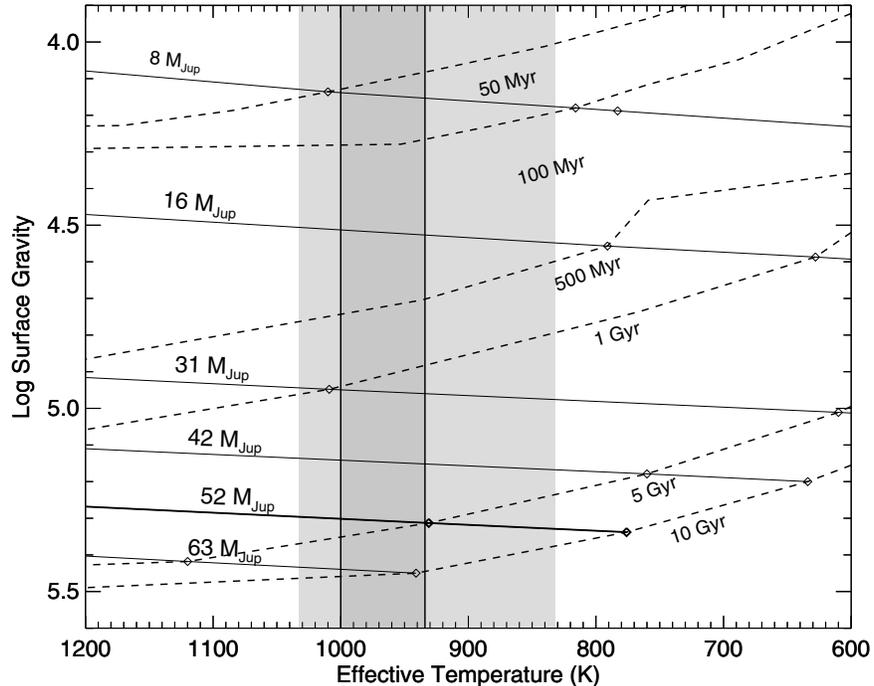
**Figure 4.** Posterior distributions using  $10^6$  MCMC steps for *JH*-band (purple), *J*-band (blue), and *H*-band (red) spectra. Histograms show the distributions marginalized over gravity (top) and temperature (left). The distributions for different bands show multiple peaks indicating that higher spectral resolution measurements could further refine our understanding of the companion physical properties. HD 19467 B has a cool temperature,  $T_{\text{eff}} = 978^{+20}_{-43}$  K, and surface gravity,  $\log g = 4.87^{+0.44}_{-0.66}$  (see text for discussion).

$\approx 0.3$  Gyr companion HN Peg B (Luhman et al. 2007). We find that the most recent BT-Settl models reproduce the  $R \approx 30$  spectra for cold dwarfs later than T5. All of these objects are  $> 1$  Gyr old with surfaces gravities predicted to be  $> 5.0$  dex (e.g., Baraffe et al. 2003), with the exception of HN Peg B. Model fits for field objects later than T5 result in best fit surface gravities  $\log g = 4.1\text{--}4.9$  dex, while the best fit surface gravity for HN Peg B, which is known to be young, is  $> 5.0$  dex. Therefore we conclude that our surface gravity results are not reliable, likely a combination of the low spectral resolution perhaps exacerbated by a systematic bias in the BT Settl synthetic spectra.

## 5. EVOLUTIONARY MODELS

Initial analysis of the HD 19467 system age from Crepp et al. (2014) yielded two distinct results: a gyro-chronological age of  $4.3^{+1.0}_{-1.2}$  Gyr, and an isochronal age of  $9 \pm 1$  Gyr. The younger age was adopted in that paper due to convergence issues when iterating the “Spectroscopy Made Easy” program for tracks near the end of the stellar evolutionary model grid (Valenti and Fischer 2005). Given the effective temperature range that we find for HD 19467 B, we can now help discriminate between the above disparate age estimates. Figure 5 shows the best-fit model parameters for surface gravity and effective temperature overplotted against COND03 isochrones and mass tracks (Baraffe et al. 2003). The gray shaded region indicates the allowable parameter space for  $T_{\text{eff}}$  discounting the (possibly spurious) surface gravity results. The dark solid line corresponds to the mass lower-limit of  $M > 52 M_{\text{Jup}}$  derived for HD 19467 B from RV measurements.

We find that the gyrochronological age of  $4.3^{+1.0}_{-1.2}$  Gyr corresponds to masses that lie near the edge of acceptable values ( $m > 52 M_{\text{Jup}}$ ) for the range of temperatures allowed by spectral



**Figure 5.** Best-fit model parameters for surface gravity and effective temperature plotted onto COND03 isochrones and mass tracks (Baraffe et al. 2003). The dark gray shaded region shows the  $1\sigma$  parameter space for the full  $JH$  spectral fit discounting the (likely spurious) surface gravity results. The light-gray shaded region indicates the  $1\sigma$  parameter space for  $J$  and  $H$  spectral fits. The effective temperature derived from spectral fits shows consistency with the  $9 \pm 1$  Gyr isochronal age of the primary star and  $M > 52 M_{\text{Jup}}$  mass constraint (solid line) derived from the RV acceleration (Crepp et al. 2014). The  $4.3^{+1.0}_{-1.2}$  Gyr gyro-chronological age is less consistent with the companion mass and effective temperature though cannot be ruled out.

fitting. The companion would require an edge-on orbit in this case. Examining the intersection of the hottest allowable temperature, i.e.,  $1\sigma$  (998 K), from  $JH$  spectral fitting and smallest allowable mass track ( $M = 52 M_{\text{Jup}}$ ), we find that ages  $\geq 4.6$  Gyr satisfy all available constraints (Baraffe et al. 2003). This result favors the  $9 \pm 1$  Gyr isochronal age, which was derived in Crepp et al. (2014) from color–magnitude diagrams and high resolution stellar spectroscopy, but does not rule out the gyro-chronological age. In the absence of any additional diagnostics, we adopt the 4.6–10 Gyr range as the new current best estimate for the age of HD 19467. While it is not clear why the estimated gyrochronological age differs significantly from the isochronal age, it is perhaps not surprising to find a slightly evolved G3V stars with sub-solar metallicity that is older than the Sun. As discussed in the introduction, the age of HD 19467 will be subject to further testing with interferometry by directly measuring the radius of the primary star.

## 6. SUMMARY AND CONCLUSIONS

We present the first direct spectrum of currently the only T dwarf companion known to produce a Doppler acceleration around a solar type star. Our observations affirm the cold nature of the companion reported in the discovery article. By self-consistently comparing IRTF SpeX observations of cold brown dwarfs to low-resolution  $JH$  spectra obtained by the P1640 IFS at Palomar, we find that HD 19467 B has an effective temperature of  $T_{\text{eff}} = 978^{+20}_{-43}$  K. The spectral-type originally assigned by Crepp et al. (2014) is confirmed; our  $R \approx 30$  observations indicate that only a marginal refinement from  $\approx T5\text{--}T7$  to  $T5.5 \pm 1$  is required.

Several lines of evidence indicate that HD 19467 B is a substellar object. P1640 spectra reveal significant methane

absorption in the  $H$ -band. At  $R \approx 30$ , we measure a spectral index of  $S_{CH_4} = 3.1^{+0.6}_{-0.8}$ , consistent with that of a late T dwarf. At a temperature of 978 K, evolutionary models predict that such an object must be substellar at any age (Baraffe et al. 2003). Further, the dynamical mass limit derived from joint RV and imaging measurements yields a value of  $M \geq 52 M_{\text{Jup}}$ , well within the brown dwarf range.

Comparing the  $T_{\text{eff}}$  of HD 19467 B to evolutionary models resulted in a necessary revision to the companion age from  $4.3^{+1.0}_{-1.2}$  Gyr (gyrochronological estimate) to 4.6–10 Gyr. This is the only way to establish a self-consistent connection between our new spectroscopy results and the companion mass lower limit derived from the Doppler trend.

HD 19467 B presents a rare opportunity to study a cold brown dwarf for which mass, age, and metallicity information is available simultaneously in addition to a direct spectrum. Other comparable objects include GJ 758 B (Thalmann et al. 2009; Janson et al. 2011) and GJ 504 b (Kuzuhara et al. 2013). These late-type companions orbit nearby, well-characterized G-stars, are older than directly imaged exoplanets, and exhibit methane absorption (Janson et al. 2011, 2013), although do not yet have RV accelerations or spectra reported in the literature. Together, they may ultimately be used as a control group to study the effects of surface gravity on the atmospheres of the oldest and coldest benchmark TY-dwarfs.

Theoretical spectral models and evolutionary models have yet to be calibrated over a wide range of temperatures and surface gravities. Indeed, the few benchmark objects that have been studied in sufficient detail to test these models have already revealed discrepancies in the luminosity values predicted from brown dwarf evolution (Dupuy et al. 2014). Unlike T dwarfs discovered as field objects, the physical properties of HD 19467 B may be measured without relying upon theoretical models.

Our current best estimates of the companion mass, age, and metallicity are:

Mass:  $\geq 51.9^{+3.6}_{-4.3} M_{\text{Jup}}$  (dynamics),

[Fe/H]:  $-0.15 \pm 0.04$  (stellar spectroscopy),

Age: 4.6–10 Gyr (multiple techniques).

respectively. The first two quantities have been determined without any reference to the emergent spectrum of light received from the companion itself. Only the age lower-limit has been inferred using brown dwarf theoretical evolutionary models. The age upper limit was determined through stellar isochrones. In time, the uncertainty in each will narrow through continued Doppler and astrometric monitoring as well as further investigations of its host star, which has a precise parallax and conveniently resembles the Sun.

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