

DETECTING THE COMPANIONS AND ELLIPSOIDAL VARIATIONS OF RS CVN PRIMARIES:
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ABSTRACT

To measure the properties of both components of the RS CVn binary σ Geminorum (σ Gem), we directly detect the faint companion, measure the orbit, obtain model-independent masses and evolutionary histories, detect ellipsoidal variations of the primary caused by the gravity of the companion, and measure gravity darkening. We detect the companion with interferometric observations obtained with the Michigan InfraRed Combiner (MIRC) at Georgia State University's Center for High Angular Resolution Astronomy (CHARA) Array with a primary-to-secondary H -band flux ratio of 270 ± 70 . A radial velocity curve of the companion was obtained with spectra from the Tillinghast Reflector Echelle Spectrograph (TRES) on the 1.5-m Tillinghast Reflector at Fred Lawrence Whipple Observatory (FLWO). We additionally use new observations from the Tennessee State University Automated Spectroscopic and Photometric Telescopes (AST and APT, respectively). From our orbit, we determine model-independent masses of the components ($M_1 = 1.28 \pm 0.07 M_\odot$, $M_2 = 0.73 \pm 0.03 M_\odot$), and estimate a system age of 5 ∓ 1 Gyr. An average of the 27-year APT light curve of σ Gem folded over the orbital period ($P = 19.6027 \pm 0.0005$ days) reveals a quasi-sinusoidal signature, which has previously been attributed to active longitudes 180° apart on the surface of σ Gem. With the component masses, diameters, and orbit, we find that the predicted light curve for ellipsoidal variations due to the primary star partially filling its Roche lobe potential matches well with the observed average light curve, offering a compelling alternative explanation to the active longitudes hypothesis. Measuring gravity darkening from the light curve gives $\beta < 0.1$, a value slightly lower than that expected from recent theory.

Subject headings: binaries: close – stars: activity – stars: imaging – stars: individual (σ Geminorum)
– stars: variables: general

1. INTRODUCTION

RS Canum Venaticorum (RS CVn) stars are spotted, active binary systems exhibiting photometric and Ca H and K variability (Hall 1976). Often tidally-locked, these systems are composed of an evolved primary star (giant or subgiant) and a subgiant or dwarf companion (Berdyugina 2005; Strassmeier 2009). With active binaries, not only is there potential to determine the component masses and system evolutionary history but also to understand the magnetic field interactions through active longitudes, particular longitudes 180° apart with persistent, long-lived starspots (Berdyugina & Tuominen 1998; Berdyugina 2005).

Observing the magnetic phenomena of rapidly-rotating evolved stars also sheds light on the magnetic activity of rapidly-rotating young stars, such as T Tauri stars. Both T Tauri and RS CVn systems have starspots analogous to sunspots—cool starspots resulting from stifled convection in the outer layers of the stars due to strong

magnetic fields (Petrov 2003; Berdyugina 2005).

σ Geminorum (σ Gem, HD 62044, HIP 37629, HR 2973) is an RS CVn system known to exhibit starspots, often ascribed to “active longitudes” (e.g., Hall et al. 1977; Henry et al. 1995). The system has been characterized as a single-lined spectroscopic binary (Herbig & Spalding 1955) with a K1III primary (Roman 1952). The orbital period of σ Gem is slightly longer than the primary star's rotation period derived from the fastest rotating spots ($P_{\text{orb}} = 19.60$ days, $P_{\text{rot,min}} = 19.47$ days; Kajatkari et al. 2014).

Because of its large starspots, σ Gem is a frequent target for understanding starspot evolution. Eberhard & Schwarzschild (1913) first reported σ Gem as active and potentially spotted due to fluctuations in the Ca H and K lines as the star rotated. Decades later, Hall et al. (1977) identified photometric variations suggesting starspots ($\Delta V \sim 0.07$). Initial models of the surface of σ Gem often showed the surface with two starspots oriented on opposite sides of the primary star (Fried et al. 1983). Berdyugina & Tuominen

(1998) emphasize that, due to tidal locking, the starspots are located such that one spot constantly faces the companion and the other spot is 180° offset. The majority of spot models applied to light curves of σ Gem consist of two spots on a spherical star (Eker 1986; Strassmeier et al. 1988; Olah et al. 1989; Henry et al. 1995; Jetsu 1996; Padmakar & Pandey 1999; Kajatkari et al. 2014). Doppler images have suggested the surface is covered with a larger number of smaller spots (Hatzes 1993; Kóvári et al. 2001, 2014).

To understand the binary system, we present our analysis of the first detections of the companion in our interferometric and radial velocity data sets, as well as photometric data. In Section 2, we describe the observations for our data sets. In Section 3, we discuss our analysis of the data sets, including the first astrometric and spectroscopic detections of the companion star and orbital parameters. In Section 4, we present evolutionary constraints and a Hertzsprung-Russell (H-R) diagram. In Section 5, we discuss our analysis of the photometric data set, including detected ellipsoidal variations and measured gravity darkening. In Section 6, we present the conclusions of our study of σ Gem.

2. OBSERVATIONS

2.1. Interferometry

We obtained interferometric data with Georgia State University’s Center for High-Angular Resolution Astronomy (CHARA) Array. The CHARA Array is a Y-shaped array of six 1-m class telescopes with non-redundant baselines varying from 34- to 331-m located at Mount Wilson Observatory, California (ten Brummelaar et al. 2005). Using all six telescopes and the Michigan InfraRed Combiner (MIRC; Monnier et al. 2004, 2006), we obtained H -band ($1.5 - 1.8 \mu\text{m}$) data (eight channels across the photometric band with $\lambda/\Delta\lambda \sim 40$) on UT 2011 Nov 9 and Dec 7, 8, 9; 2012 Nov 7, 8, 21, 22, 24, 25 and Dec 4, 5.

We made detections of the companion in the data from UT 2011 Dec 8; 2012 Nov 7, 8, 24, and 25. The remaining nights of observation had insufficient uv coverage due to poor seeing or short observation lengths, leaving the companion undetected. We reduced and calibrated these data with the standard MIRC pipeline (see Monnier et al. 2007, 2012; Zhao et al. 2009; Che et al. 2011, for pipeline details). We used at least one calibration star for each night of data (see Table 1).

2.2. Radial Velocity

To constrain the spectroscopic orbit for σ Gem, we utilized three independent sets of radial-velocity data: two sets of single-lined velocities for the primary, and a new set of double-lined velocities for both components of the binary.

One set of the radial velocity measurements for the primary star was published in Massarotti et al. (2008). These 39 data points were obtained with two identical CfA Digital Speedometers (Latham 1992) on the 1.5-m Wyeth Reflector (Oak Ridge Observatory) and 1.5-m Tillinghast Reflector (Fred Lawrence Whipple Observatory) telescopes (2003 December 30 – 2007 June 5).

From 2012 October 1 – 2015 January 9, using the Tillinghast telescope with the Tillinghast Reflector Echelle

Spectrograph (TRES; Fűrész 2008), we were able to make fifteen detections of the secondary spectra for the first time. Along with sixteen new primary star measurements, these new radial velocities are presented in Table 2. We add 0.14 km s^{-1} to these sets of radial velocities to account for these data being reported on the CfA native system (Stefanik et al. 1999, note the correction is inaccurately stated as a subtraction in this reference). For details of these observations and data analysis, see Appendix A.

The additional radial velocity data set consists of 43 spectrograms of the primary star of σ Gem taken between 2009 January 12 – 2014 December 1 with the Tennessee State University 2-m automatic spectroscopic telescope (AST), fiber-fed echelle spectrograph, and a CCD detector at Fairborn Observatory, Arizona (see Table 3; Eaton & Williamson 2004, 2007). At first, the detector was a 2048×4096 SITE ST-002A CCD with $15 \mu\text{m}$ pixels. Eaton & Williamson (2007) discussed the reduction of the raw spectra and wavelength calibration. Those echelle spectrograms have 21 orders that cover the wavelength range $4920\text{--}7100 \text{ \AA}$ with an average resolution of 0.17 \AA , corresponding to a resolving power of 35000 at 6000 \AA . Those spectra have a typical signal-to-noise value of 30.

In the summer of 2011 the AST SITE CCD and its dewar were retired and replaced with a Fairchild 486 CCD, a 4096×4096 array of $15 \mu\text{m}$ pixel, that is housed in a new dewar. With the new CCD the wavelength coverage ranged from 3800 to 8600 \AA . The resolution was reduced slightly to 0.24 \AA or a resolving power of 25000 at 6000 \AA . These more recent spectra have signal-to-noise ratios of about 70.

Fekel et al. (2009) provided an extensive general description of velocity measurement of the Fairborn AST spectra. In the case of σ Gem, we measured a subset of 63 lines from our solar-type star line list that covers the $4920\text{--}7120 \text{ \AA}$ region. Because the lines of σ Gem have significant rotational broadening, we fit the individual lines with a rotational broadening function. The Fairborn velocities are on an absolute scale. A comparison of our unpublished measurements of several IAU standard stars with those determined by Scarfe et al. (1990) indicates that the Fairborn Observatory velocities from the SITE CCD have a small zero-point offset of -0.3 km s^{-1} . Velocities from the Fairchild CCD spectra have a slightly larger zero-point offset of -0.6 km s^{-1} relative to those of Scarfe et al. (1990). Thus, in Table 3 we corrected our measured velocities by either 0.3 or 0.6 km s^{-1} , depending on which detector was used.

2.3. Photometry

We used differential photometry of σ Gem and a comparison star from the Tennessee State University T3 0.4-m Automated Photometric Telescope (APT) located at Fairborn Observatory, Arizona. For details on the observational procedure and photometers see Henry (1999) and Fekel et al. (2005).

The differential Johnson B and V light curves cover 1987 November 21 – 2015 March 13 (see Table 4 and Figure 1). Subsets of these data were analyzed by Henry et al. (1995) and Kajatkari et al. (2014). For the first time, we make the full set of T3 APT photometry

TABLE 1
 CALIBRATORS FOR σ GEMINORUM

Calibrator Name	Calibrator Size (mas)	Source	UT Date of Observation
HD 37329	0.71 ± 0.05	Bonneau et al. (2006)	2012 Nov 8
HD 50019 (θ Gem)	0.81 ± 0.06	Bonneau et al. (2006)	2012 Nov 7, 8, 25
HD 63138	0.65 ± 0.04	MIRC calibration	2011 Dec 8; 2012 Nov 8
HD 69897 (χ Cnc)	0.73 ± 0.05	Bonneau et al. (2006)	2012 Nov 7, 24, 25

 TABLE 2
 RADIAL VELOCITY DATA OF σ GEM (CfA)

HJD -2400000	Primary (km s ⁻¹)	Secondary (km s ⁻¹)
56202.0199	10.66	103.44
56230.0436	77.98	-9.80
57002.9567	14.17	94.68
57003.9301	9.71	101.50
57014.8600	79.02	-12.11
57015.8427	75.96	-10.77
57018.8519	51.08	28.91
57019.9181	38.76	60.57
57020.9452	27.84	78.02
57021.8962	19.05	84.20
57024.9812	9.10	104.10
57025.9317	12.59	99.41
57026.8891	18.88	85.69
57028.9057	38.58	
57029.8670	49.81	27.28
57031.8806	69.76	-2.88

NOTE. — Errors on the primary radial velocities are 0.84 km s^{-1} . Errors on the secondary radial velocities are 3.8 km s^{-1} . These were then scaled for our orbit fit to have a total $\chi^2 = 1.00$. Note these radial velocities are on the native CfA system. We added 0.14 km s^{-1} for use in our analysis (Stefanik et al. 1999).

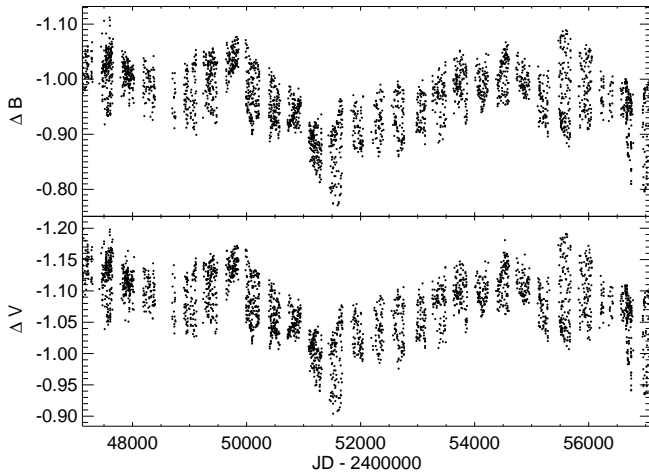


FIG. 1.— Johnson B and V differential magnitudes of σ Gem acquired over 28 observing seasons from 1987 – 2015 with the T3 0.4-meter APT at Fairborn Observatory in southern Arizona.

of σ Gem available in electronic format in Table 4.

3. ORBITAL ELEMENTS

In order to derive the astrometric orbit of σ Gem, we searched for the companion with model fitting. We modeled the system with the resolved primary star and an unresolved secondary. We allowed the primary radius along the major axis, primary major-to-minor axis ratio, primary major axis position angle, primary-to-secondary flux ratio, and secondary position to vary. During the fit-

 TABLE 3
 RADIAL VELOCITY DATA OF σ GEM
 (AST/TSU)

HJD -2400000	Primary (km s ⁻¹)
54843.9716	38.0
54844.8597	28.4
54845.6849	20.5
54846.6455	14.0
54847.6848	9.2
54848.7712	9.2
54849.7451	12.1
54850.6645	17.6
54856.9772	75.7
54859.7922	73.1
54860.6910	66.7
54861.6447	58.7
54862.6829	48.3
54863.6434	37.3
54864.6443	26.7
54865.6596	17.8
54866.6401	11.7
54867.6456	9.1
54868.9572	10.3
54869.6594	13.8
54870.6346	21.1
54975.6558	77.9
54976.6372	76.4
54981.6497	33.3
54982.6541	23.0
54983.6544	15.2
54984.6546	9.9
54985.6548	8.5
54986.6547	10.7
55060.9627	23.5
55061.9642	16.0
55066.9516	23.3
55069.9622	54.3
55070.9596	64.0
55071.9572	71.9
55072.9375	76.6
56984.7820	9.4
56987.7638	19.7
56988.7240	28.0
56989.7317	38.9
56990.7354	50.0
56991.7355	60.9
56992.8661	70.8

NOTE. — Errors on the primary radial velocities are 0.3 km s^{-1} . These data were then scaled for our orbit fit to have a total $\chi^2 = 1.00$.

ting, we weighted the data such that the separate observables (squared visibilities, closure phases, and triple amplitudes) contributed to the final χ^2 with equal weight. The parameter errors for the primary star size and the primary-to-secondary flux ratio were based on the epoch-to-epoch variation, while the relative positional error of the secondary compared to the primary were based on the residuals to the orbit fit (see discussion on orbit fitting).

The coordinates of the detections on five nights (UT

TABLE 4
JOHNSON B AND V DIFFERENTIAL MAGNITUDES OF σ GEM FROM THE T3 APT

HJD -2400000	ΔB (σ Gem - HD 60318 ^a)	ΔV (σ Gem - HD 60318)	ΔB (v Gem ^b - HD 60318)	ΔV (v Gem - HD 60318)
47121.0481	-1.006	-1.123		
47122.0465	-1.010	-1.135		
47125.0364	-1.030	-1.125	-0.766	-1.303
47126.0364	-1.011	-1.128	-0.757	
47128.0310	-0.975		-0.759	

NOTE. — Table 4 is published in its entirety in the electronic edition. A portion is shown here for guidance regarding its form and content.

^a HD 60318, $V = 5.33$, $B - V = 1.01$

^b v Gem (HD 60522, 69 Gem), $V = 4.06$, $B - V = 1.54$

2011 Dec 8; 2012 Nov 7, 8, 24, and 25 are listed in Table 5). The H -band flux ratio for the primary star to the secondary is 270 ± 70 . In addition to detecting the secondary star, we measured the uniform disk diameter of the primary to be $\theta_{UD,1} = 2.335 \pm 0.007$ mas (limb-darkened disk diameter $\theta_{LD,1} = 2.417 \pm 0.007$ mas) with a major-to-minor axis ratio of 1.02 ± 0.03 . Our measurements are slightly larger than those in the CHARM2 catalog (uniform disk diameter of $\theta_{UD,1} = 2.18 \pm 0.05$ mas, limb-darkened disk diameter of $\theta_{LD,1} = 2.31 \pm 0.05$ mas; Richichi et al. 2005).

To determine the binary orbit, we simultaneously fit our interferometric and radial velocity data with Monte Carlo realizations. The five interferometric points are as described above, and we present the scaled error bars of the major and minor axis in Table 5 to give our fit a total $\chi^2 = 1.00$. For the radial velocity data we combine the Massarotti et al. (2008, adding 0.14 km s^{-1} to account for the values reported on the CfA native system), new CfA data, and the AST data to fit simultaneously with the astrometry. The radial velocity errors are similarly scaled ($\text{rms}_{\text{CfA},1} = 0.84 \text{ km s}^{-1}$, $\text{rms}_{\text{AST},1} = 0.3 \text{ km s}^{-1}$, $\text{rms}_{\text{CfA},2} = 3.8 \text{ km s}^{-1}$).

Using the complete radial velocity data sets, we find an eccentricity of $e = 0.014 \pm 0.004$, consistent with slightly eccentric orbits reported by Harper (1935), Pourbaix et al. (2004), and Massarotti et al. (2008). However, Luyten (1936), Batten et al. (1978), and Dümmler et al. (1997) reported a circular orbit. To investigate this discrepancy, we used the APT light curve to eliminate the primary star’s radial velocity data that were obtained when σ Gem presented starspots ($\Delta V > 0.04$), as these could cause shifts in the velocities (e.g., Saar & Donahue 1997). The remaining primary star radial velocity data obtained when σ Gem did not exhibit large starspots from the Massarotti et al. (2008)/CfA data set span 2006 December 6 – 2007 June 5, and those from the AST data set span 2009 January 12 – Jun 4. Using the primary star’s truncated data set with 42% of the Massarotti et al. (2008)/CfA and 33% of the AST epochs removed, we find the orbit is consistent with a circular orbit, $e = 0.002 \pm 0.002$, and we adopt a circular orbit for the rest of this paper.

Requiring eccentricity $e = 0$ and the argument of periastron for the primary $\omega = 0^\circ$ the simultaneous Monte Carlo realizations gave the orbital parameters and their $1 - \sigma$ errors listed in Table 6. The visual orbit is illustrated in Figure 2, and the radial velocity curve is presented in Figure 3. We use the conventions presented by Heintz (1978), where the argument of periastron, ω , and the time of nodal passage (maximum recessional ve-

locity), T_0 , are defined by the primary star’s orbit. The ascending node, Ω , is independent of definition, being equivalent with respect to either the primary or secondary star.

Our orbital parallax, $\pi = 25.8 \pm 0.4$ mas can be compared with the Hipparcos parallax of 26.68 ± 0.79 mas (ESA 1997). As an unresolved binary with a variable component, σ Gem does not exhibit the photocenter shifts found to be troublesome for measuring binary system parallax with Hipparcos (ESA 1997; Halbwachs & Pourbaix 2005). Assuming that the secondary is negligibly bright, the semi-major axis of the photocentric orbit of the primary is at most 1.71 mas wide, which is at the limit of detectability (Pourbaix 2002) for Hipparcos. Combining Hipparcos data and our visual orbit, the parallax is 26.4 ± 0.8 mas, consistent with our orbital parallax. For our subsequent analysis, we adopt our higher-precision orbital parallax, $\pi = 25.8 \pm 0.4$ mas.

With a circular orbit and $P_{\text{orb}} \sim P_{\text{rot}}$ (e.g., Kajatkari et al. 2014), we expect σ Gem to have aligned rotational and orbital axes. Given our orbital and stellar parameters, we can calculate the obliquity of the system. Comparing our calculated value of $v \sin i = (2\pi R_1 / P_{\text{orb}}) \times v \sin i = 24.8 \pm 0.4 \text{ km s}^{-1}$ with the observational rotational velocity of $v \sin i = 26.7 \pm 0.5 \text{ km s}^{-1}$ (from the TRES spectra), we find that the calculation is smaller than the observational value. This discrepancy could be attributed to the estimate of microturbulence or the presence of the large spot structures on the surface of σ Gem during the TRES observations instead of a small, non-zero obliquity.

4. MASSES AND HERTZSPRUNG-RUSSELL DIAGRAM

Using our complete orbital fit, we obtain model-independent masses $M_1 = 1.28 \pm 0.07 M_\odot$ and $M_2 = 0.73 \pm 0.03 M_\odot$. With the stellar parameters of the primary star (including $T_{\text{eff},1} = 4530 \pm 60$ K, see Table 6) and the primary-to-secondary H -band flux ratio detected using the CHARA/MIRC data (270 ± 70), we are able to constrain the parameters (luminosity, temperature, and radius) of the secondary star. We use the flux ratio and NextGen stellar atmospheres (Hauschildt et al. 1999) to constrain the stellar flux to calculate a range of luminosities for reasonable effective temperatures (4000–4700 K) for a $0.73 \pm 0.03 M_\odot$ main sequence star (see Figure 4). We obtain a range of luminosities ($0.11 - 0.15 L_\odot$) and radii ($0.70 - 0.59 R_\odot$). We note that our analysis predicts a primary-to-secondary Johnson V -band flux ratio of 290 assuming $T_{\text{eff},2} = 4500$ K), which is not in agreement with the flux ratio given by the spectroscopic 519 nm

TABLE 5
 DETECTIONS FOR THE COMPANION OF σ GEMINORUM WITH RESPECT TO THE PRIMARY

UT Date	JD - 2400000	Separation (mas)	Position Angle ($^\circ$) ^a	Error Ellipse Major Axis (mas) ^b	Error Ellipse Minor Axis (mas) ^b	Error Ellipse Position Angle ($^\circ$) ^a	Reduced χ^2
2011 December 08	55903.95	2.83	19.1	0.30	0.09	80	4.4
2012 November 07	56238.97	4.32	8.6	0.04	0.03	280	2.8
2012 November 08	56239.86	4.68	359.7	0.13	0.06	300	2.1
2012 November 24	56256.00	2.03	39.0	0.08	0.06	30	1.5
2012 November 25	56256.95	3.12	21.6	0.05	0.04	320	1.7

NOTE. — These detections give an H -band ($1.5 - 1.8 \mu\text{m}$) flux ratio for σ Gem primary to secondary of 270 ± 70 . The uniform disk fit for the primary star is $\theta_{\text{UD},1} = 2.335 \pm 0.006$ mas (limb-darkened disk diameter $\theta_{\text{LD},1} = 2.417 \pm 0.006$ mas) with a 1.02 ± 0.03 major-to-minor axis ratio.

^aEast of North

^bScaled error bars to ensure a total $\chi^2 = 1.00$ as described in Section 3.2.

 TABLE 6
 ORBITAL AND STELLAR PARAMETERS OF σ GEM

Measured Parameters	Value
semi-major axis, a (mas)	4.63 ± 0.04
eccentricity, e	0
inclination, i ($^\circ$)	107.7 ± 0.8
argument of periastron, ω ($^\circ$) ^a	0
ascending node, Ω ($^\circ$)	1.2 ± 0.8
period, P_{orb} (days)	19.6027 ± 0.0005
time of nodal passage, T_0 (HJD) ^b	2453583.98 ± 0.03
velocity semi-amplitude, K_1 (km s^{-1})	34.62 ± 0.08
velocity semi-amplitude, K_2 (km s^{-1})	60 ± 2
system velocity, γ (km s^{-1})	43.41 ± 0.08
uniform disk diameter, $\theta_{\text{UD},1}$ (mas)	2.335 ± 0.007
limb-darkened disk diameter, $\theta_{\text{LD},1}$ (mas) ^c	2.417 ± 0.007
primary major-to-minor axis ratio	1.02 ± 0.03
H -band flux ratio, primary to secondary	270 ± 70
orbital parallax, π (mas)	25.8 ± 0.4
distance, d (pc)	38.8 ± 0.6
Derived Parameters	
average primary radius, R_1 (R_\odot) ^d	10.1 ± 0.4
primary luminosity, L_1 (L_\odot)	39 ± 2
primary surface gravity, $\log g_1$ (cm/s^2)	2.54 ± 0.02
primary mass, M_1 (M_\odot)	1.28 ± 0.07
secondary mass, M_2 (M_\odot)	0.73 ± 0.03
system age (Gyr)	5 ∓ 1
Literature Parameters	
primary effective temperature, $T_{\text{eff},1}$ (K) ^e	4530 ± 60
primary metallicity (iron), Fe/H ^f	0.0

NOTE. — ^aRadial velocity convention for primary with respect to the center of mass.

^bTime of maximum recessional velocity of the primary star.

^cWe applied a 3.5% correction from uniform to limb-darkened disk diameter. This is equivalent to a limb-darkening coefficient $\alpha = 0.27$.

^dUsing limb-darkened disk diameter.

^e $T_{\text{eff},1}$ is an average of temperatures given by Głęboczi & Stawikowski (1979); Poe & Eaton (1985); Stawikowski & Głęboczi (1994); O’Neal et al. (1996); Kóvári et al. (2001); Massarotti et al. (2008). The $1 - \sigma$ error is the standard deviation of these values.

^f $[\text{Fe}/\text{H}] = -0.02$ (Mallik 1998); approximated as $[\text{Fe}/\text{H}] = 0.00$.

light ratio (~ 70 primary-to-secondary; see Appendix A for details on this measurement). In order for our flux ratios to be in agreement, the secondary star would have $T_{\text{eff},2} = 6400$ K, which is not consistent with the spectroscopic observations, nor with a main-sequence star given the location on the H-R diagram. We cannot rule out the effect of starspots on the discrepant flux ratios as these were not accounted for when interferometrically detecting the companion and the spot features present during the interferometric and spectroscopic observations differ as evidenced in the APT light curve. Additionally, Prato et al. (2002) and Lehmann et al. (2013) also reported discrepancies between TODCOR-reported flux

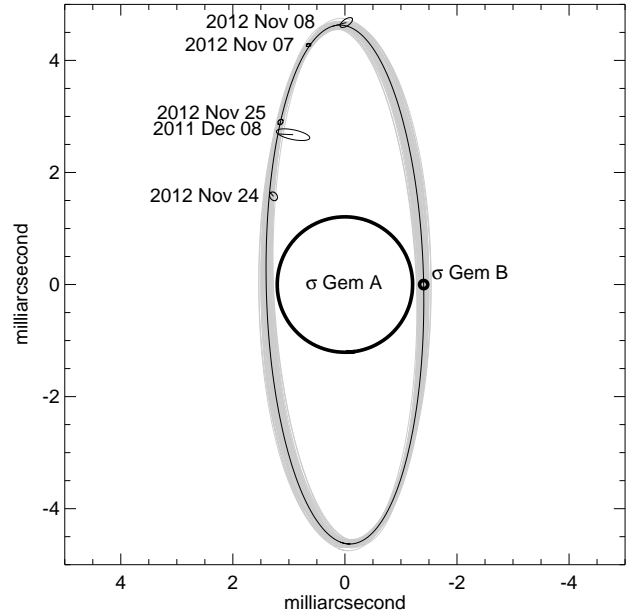


FIG. 2.— Visual orbit for the prototypical RS CVn system σ Gem with our observed stellar primary radius (thick black line, σ Gem A) and our dates of companion detection and their locations on the orbit (black error ellipses). The predicted radius of the companion star, σ Gem B, is plotted for scale with the small thick black circle. The orbits of fifty Monte Carlo realizations are presented as the light gray orbits. Black lines connect the center of the detection error ellipse to the expected point in the best-fit orbit, which is overlaid in black (given in Table 6 with $1 - \sigma$ errors). At the southernmost point in the orbit, the secondary star is moving toward the observer. Note: axis units are milliarcsseconds (mas) with north upwards and east to the left.

ratios and their expected values. Therefore, we use only the H -band flux ratio.

We plot the location of the components of σ Gem on an H-R diagram, as well as the corresponding evolutionary tracks. We use Dartmouth stellar evolution tracks ($\text{Fe}/\text{H} = 0.0$, $\alpha/\text{Fe} = 0.0$, PHOENIX-based models; Dotter et al. 2008) for the interpolated model masses ($M_{1,\text{model}} = 1.28 \pm 0.07 M_\odot$, $M_{2,\text{model}} = 0.73 \pm 0.03 M_\odot$). Our primary falls nearly on the $1.28 M_\odot$ evolutionary track with an estimated temperature of 4530 ± 60 K (Głęboczi & Stawikowski 1979; Poe & Eaton 1985; Stawikowski & Głęboczi 1994; O’Neal et al. 1996; Kóvári et al. 2001; Massarotti et al. 2008). The range of locations for the secondary on the H-R diagram passes

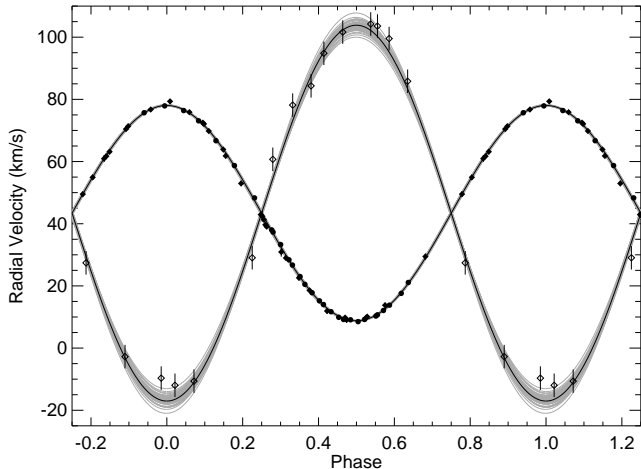


FIG. 3.— Radial velocity curves of σ Gem. The filled diamonds represent our sample of measured observations from Massarotti et al. (2008)/CfA, and the filled circles are the AST observations. Both data sets are restricted to those data points obtained with no starspots present (see Section 3). $1 - \sigma$ errors in velocity are presented unless the error is smaller than the diamonds and circles. The radial velocity curves of fifty Monte Carlo realizations are presented as the light gray orbits. The radial velocity for the best orbital parameters is overlaid in black. Similarly, the open diamonds represent CfA radial velocity data for the secondary star with $1 - \sigma$ error bars. The light gray orbits are fifty Monte Carlo realizations with the best orbital parameters overlaid in black. See Table 6 for orbital parameters with $1 - \sigma$ errors.

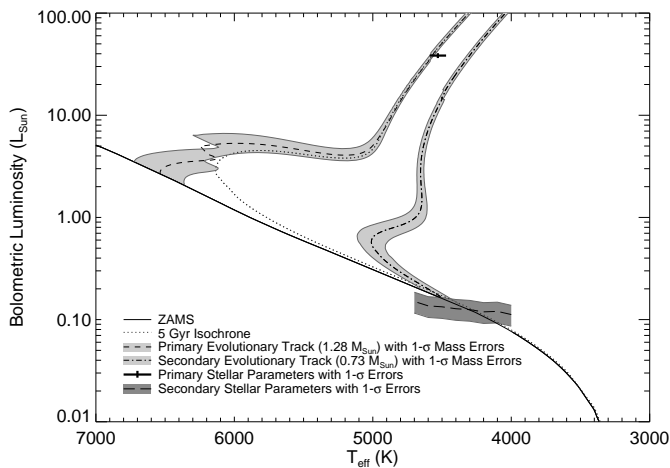


FIG. 4.— H-R diagram for σ Gem. The dashed and dot-dashed lines are the main sequence and post-main sequence evolutionary tracks for $1.28 M_{\odot}$ and $0.73 M_{\odot}$ stars with $[Fe/H] \sim 0.0$, respectively (Dotter et al. 2008). The gray regions represent our $1 - \sigma$ mass errors ($M_1 = 1.28 \pm 0.07 M_{\odot}$; $M_2 = 0.73 \pm 0.03 M_{\odot}$) with the solid black line representing the zero age main sequence. The dotted line is a 5 Gyr isochrone (PHOENIX; Dotter et al. 2008). The measured location of the primary with $1 - \sigma$ errors is indicated by the plus sign. The region where the companion could be located given our flux ratio and reasonable temperature estimates is indicated with the long-dashed line (with $1 - \sigma$ errors in luminosity).

through the main sequence for a star of $0.73 M_{\odot}$. We find an age of the system of 5 ∓ 1 Gyr. Based upon the masses and age of the stars, we suggest that the primary star is an evolved late F-type star that is now a K giant. The secondary star is a main-sequence early K star.

5. ELLIPSOIDAL VARIATIONS AND GRAVITY DARKENING

Henry et al. (1995) and Kajatkari et al. (2014) previously published subsets of the APT light curve data for starspot modeling and measuring differential rotation. Both studies emphasized the presence of active longitudes on opposite sides of σ Gem to explain the quasi-sinusoidal variation appearing at half of the orbital period.

We removed long-term trends, folded the APT photometry over the orbital period ($P_{\text{orb}} = 19.6027$ days), and binned the data (0.025 in phase). The resultant Johnson B and V light curves are presented in Figure 5. The quasi-sinusoidal trend observed in the averaged light curves suggests the possibility of ellipsoidal variations due to distortions of the primary star partially filling its Roche lobe potential. With a Roche lobe radius of $16.5 R_{\odot}$, we obtain $R_1/R_L = 0.61$ (Eggleton 1983).

We used the light-curve-fitting software package Eclipsing Light Curve (ELC; Orosz & Hauschildt 2000) to model the ellipsoidal variations using our orbital parameters with no free parameters (gravity darkening assumed to be $\beta = 0.08$; Lucy 1967, see Figure 5). The characteristics of the ellipsoidal variations with this model as compared to the light curve of σ Gem indicate that the long-term signature likely is indeed due to ellipsoidal variations, in contrast to previous suggestions that the periodicity at $P_{\text{orb}}/2$ is due to active longitudes aligned with the orbit (e.g., Henry et al. 1995; Jetsu 1996; Berdyugina & Tuominen 1998; Kajatkari et al. 2014; Kóvári et al. 2014). We note that rotation periods derived from the analysis of the light curve (e.g., Kajatkari et al. 2014) suggest the star is rotating slightly faster than the orbital period, further supporting our identification of ellipsoidal variations in σ Gem. It should be noted that removing the effect of ellipsoidal variations from the light curve does not eliminate all starspot signatures (See Appendix B).

The ELC model fit of ellipsoidal variations can be improved to better match our data. We modeled the system again with no free parameters except for the gravity darkening coefficient, β for $T_{\text{eff}} \propto g^{\beta}$ (von Zeipel 1924), as Espinosa Lara & Rieutord (2012) recently suggested $\beta \sim 0.21$ for convective stars, substantially higher than the canonical $\beta \sim 0.08$ (Lucy 1967) value assumed in our fixed-parameter fit. Although our average light curve is still contaminated by some residual spot modulation, we find that $\beta = 0.02 \pm 0.02$ with error bars determined by bootstrapping over observing seasons of the 27 years of observation in the APT light curve. This value strongly rules out $\beta > 0.1$ for this system (see Figure 5).

6. CONCLUSIONS

In this work, we have made the first visual detections of the secondary star of σ Gem using interferometric and spectroscopic observations. We establish the first visual orbit by combining the interferometric detections with radial velocity data. The determination of orbital pa-

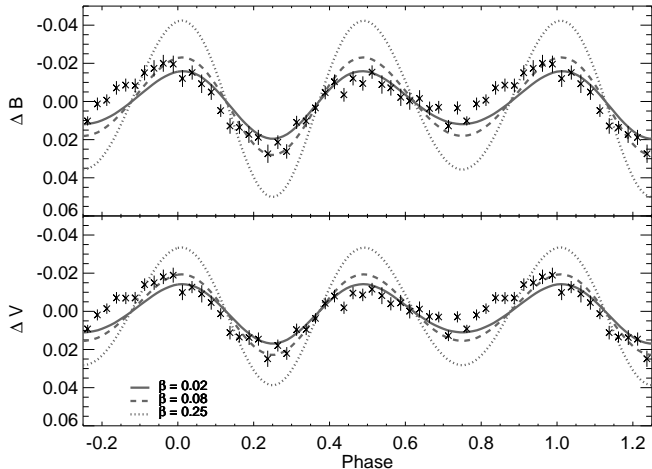


FIG. 5.— Differential folded and binned light curves of σ Gem for B and V magnitudes plotted with error bars from the binning. Each data point is an average of data points spanning 0.025 in phase from the complete light curve folded on the orbital period. The quasi-sinusoidal signature of the averaged light curve is due to ellipsoidal variations caused by the primary star partially filling its Roche lobe potential. The lines represent the ELC models for ellipsoidal variations with the gravity darkening coefficient $\beta = 0.02, 0.08$, and 0.25 , where $\beta = 0.02 \pm 0.02$ is the best fit to the binned and averaged light curves.

rameters has allowed for model-independent mass determinations ($M_1 = 1.28 \pm 0.07 M_\odot$, $M_2 = 0.73 \pm 0.03 M_\odot$).

Folded and binned photometric data have shown evidence of ellipsoidal variations, gravitational distortions of the primary star caused by the close companion. The light curve is comparable to light curve models created only from stellar and orbital parameters (assuming no starspots). Although the ellipsoidal variations are only a small effect, the primary star of σ Gem is not spherical, partially filling its Roche lobe potential and having a surface temperature gradient. Our establishment of ellipsoidal variations offers a compelling alternative explanation to the previously purported detections of active longitudes, starspots on either side of the primary star (Henry et al. 1995; Jetsu 1996; Berdyugina & Tuominen 1998; Kajatkari et al. 2014; Kóvári et al. 2014).

APPENDIX

A. CFA RADIAL VELOCITIES

The CfA radial velocity data were obtained using two identical CfA Digital Speedometers (Latham 1992) on two different telescopes: the 1.5-m Wyeth Reflector at the Oak Ridge Observatory located in the town of Harvard, Massachusetts, and the 1.5-m Tillinghast Reflector at the Fred Lawrence Whipple Observatory (FLWO) on Mount Hopkins, Arizona. In the first set of radial velocity measurements, altogether 78 observations were obtained; 32 with the Tillinghast Reflector and 46 with the Wyeth Reflector. Those radial velocities have all been published (Massarotti et al. 2008), so the details of the procedures and reductions will not be repeated here, except to provide some overall characteristics of the data. Forty of the observations were obtained over a span of less than an hour using the Wyeth Reflector. They have been averaged to give one data point for the analysis in this paper. However, those forty observations provide an opportunity to evaluate the precision of the individual velocity measurements from the CfA Digital Speedometers for an RS CVn primary with line broadening corresponding to a rotational velocity of 25 km s^{-1} . The standard deviation of a single velocity from the average of all 40 is 0.40 km s^{-1} . This compares favorably with the value of 0.45 km s^{-1} reported by Massarotti et al. (2008) for the RMS velocity residuals from their orbital solution using all 78 velocities.

To avoid confusion, the velocities determined with the CfA Digital Speedometers have always been published on the native system of the instruments. For the analysis in this paper we added 0.14 km s^{-1} to the published velocities, to

Our new orbital elements along with the folded light curve also allow for measurements of gravity darkening. We find that $\beta = 0.02 \pm 0.02$, a value of gravity darkening lower than suggested by theory (Lucy 1967; Espinosa Lara & Rieutord 2011, 2012).

In this paper, we have demonstrated that precision interferometry at CHARA is now capable of detecting the faint main-sequence companions of bright RS CVn primary stars. We are currently processing new data for other close, bright RS CVn systems and will be publishing these results in a series of follow-up papers.

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¹ Available at <http://www.jmmc.fr/searchcal>

² Available at <http://cdsweb.u-strasbg.fr/>

put them on an absolute system defined by the IAU Radial-Velocity Standard Stars (Stefanik et al. 1999, note that the sign of the correction was given as minus by mistake in that paper).

A new set of velocities for both components of the σ Gem system was obtained with the Tillinghast Reflector Echelle Spectrograph (TRES; Fűrész 2008) at FLWO during the period 2012 October 1 to 2015 January 9. TRES is a modern fiber-fed CCD echelle spectrograph with resolution of 6.7 km s^{-1} , very similar to that of the CfA Digital Speedometers. However, the free spectral range of the echelle order centered at 519 nm near the Mg b features is 10 nm, compared to 4.5 nm for the CfA Digital Speedometers. For both instruments this is the wavelength window used for the determination of absolute velocities and for TODCOR analyses. Furthermore the signal-to-noise ratio per resolution element (SNRe) achievable with the TRES CCD detector is much higher than was possible with the intensified photon-counting detectors used in the CfA Digital Speedometers. The SNRe values for the old CfA spectra ranged from 30 to 100, while the typical value for the new TRES observations is 500.

Fourteen strong TRES spectra were obtained of σ Gem over a period 30 nights in 2014 December and 2015 January with the goal of detecting the lines of the secondary and deriving a double-lined spectroscopic orbit for the first time. Fortunately, two earlier observations from 2012 October were available in the TRES archive, which provided a two-year baseline for determining a more accurate orbital period. All 16 TRES spectra were analyzed using TODCOR (Zucker & Mazeh 1994) as implemented at CfA by G. Torres, and using the CfA library of synthetic spectra to choose the optimum templates. Only one of the observations failed to give a reliable velocity for the secondary, due to close blending of the lines from the two stars.

The light of the primary dominates the composite spectrum of σ Gem, so it was straightforward to choose the template that gave the highest value for the average peak of the one-dimensional correlations; the parameters for the best template were $T_{\text{eff},1} = 4500 \text{ K}$, $\log g_1 = 2.5$ (cgs), $v \sin i_1 = 25 \text{ km s}^{-1}$, and solar metallicity. Because the secondary is so much fainter, its template parameters were only weakly constrained by the two-dimensional correlations. We therefore adopted the mass derived in this paper to guide the choice of secondary template parameters, $T_{\text{eff},2} = 4250 \text{ K}$, $\log g_2 = 4.5$ (cgs), $v \sin i_2 = 2 \text{ km s}^{-1}$, and solar metallicity; the rotational velocity was selected under the assumption of tidal synchronization of the secondary spin with the orbital period. We also tried TODCOR solutions using the nearest neighbors from our library for the secondary template, and found that the results were not sensitive to the secondary template parameters.

TODCOR has a mode in which the light ratio between the secondary and the primary in the observed wavelength window can be treated as a free parameter and thus can be determined from the spectra. We selected 10 observations where the velocities of the primary and secondary were well separated and used these to determine a light ratio at 519 nm of 0.0139 ± 0.0025 (standard deviation of the values from the mean). If the errors are well behaved in this analysis, the uncertainty of the mean light ratio from TODCOR could be formally better by about a factor of 3, or about 6%. For the final TODCOR analysis we fixed the light ratio to 0.0139 for all the observations. The TODCOR velocities and times of observation are reported in Table 2 on the same zero point as the native CfA Digital Speedometer system. For the analysis in this paper, 0.14 km s^{-1} was added to the velocities in Table 2 to put them on the IAU system.

The orbital parameters are reported in Table 7 for two different double-lined spectroscopic orbits based on just the TODCOR velocities from the 16 TRES spectra. For the first solution the eccentricity was allowed to be a free parameter. Since the derived eccentricity was not significant, $e = 0.0018 \pm 0.0029$, we also derived a solution for a circular orbit. Note that the mass ratio is well constrained; $q = 0.582 \pm 0.016$ (2.7%). Table 7 also contains the orbit reported by Massarotti et al. (2008) and the orbit fit with only the data from the AST/TSU data set. Like the results from the TRES spectra, the eccentricity from the AST radial velocities, $e = 0.002 \pm 0.003$, is consistent with a circular orbit. Note that these are the entire data sets collected at these facilities and contain radial velocity measurements that were obtained when starspots were present on the stellar surface.

Our double-lined orbital solution is among the most extreme that we have derived using the Mg b region in terms of the light ratio. Indeed, attempts to get good solutions from the neighboring echelle orders on either side of the Mg b order were unsatisfactory. Therefore we decided it would be prudent to test the reliability of our TODCOR analysis by creating sets of simulated observations using our library of synthetic spectra. For the time of each observed spectrum we simulated that observation by shifting the two synthetic template spectra by the velocities from the orbital solution and coadding after scaling by the light ratio found by TODCOR from the real observations. We found that a TODCOR analysis of the simulated observations reproduced the mass ratio and light ratio of the real data well within the estimated errors, and that the mass ratio was not sensitive to the value chosen for the light ratio. Thus we have no evidence for a systematic error in our determination of a light ratio of 0.0139 at 519 nm. This is quite different from the light ratio reported for the interferometric observations of 0.0038 at H -band.

B. DIFFERENCE LIGHT CURVES

In order to better justify our conclusion that ellipsoidal variations can explain previous claims of “active longitudes” on σ Gem, we have re-plotted some photometry from Kajatkari et al. (2014) in Figure 6 along with our prediction of the expected ellipsoidal variation component using the ELC software and system parameters from Table 6 using gravity darkening parameter $\beta = 0.02$. In Figure 6, we include data from two epochs, one showing very little overall variability and one showing high variability. In the first epoch (“Segment 8, Set 45” of Kajatkari et al. (2014)), the photometric data showed clearly a double-peaked light curve when phased with the orbital period, previously interpreted as due to active longitudes (see Kajatkari et al. 2014). Here, we now see by removing the expected ellipsoidal variation, the

TABLE 7
 σ GEM ORBITAL PARAMETERS FROM RADIAL VELOCITY CURVES^a

Parameter	TRES/TODCOR <i>e</i> free	TRES/TODCOR <i>e</i> fixed	Massarotti et al. (2008) <i>e</i> free	AST/TSU <i>e</i> free	AST/TSU <i>e</i> fixed
orbital period, P_{orb} (days)	19.6059 ± 0.0020	19.6065 ± 0.0018	19.60437 ± 0.00053	19.6041 ± 0.0002	19.6041 ± 0.0002
center of mass velocity, γ (km s^{-1})	43.554 ± 0.077	43.553 ± 0.074	43.043 ± 0.066	43.25 ± 0.07	43.33 ± 0.06
semi-amplitude, primary, K_1 (km s^{-1})	35.19 ± 0.10	35.192 ± 0.098	34.776 ± 0.100	34.60 ± 0.08	34.60 ± 0.08
semi-amplitude, secondary, K_2 (km s^{-1})	60.5 ± 1.6	60.5 ± 1.5			
eccentricity, e	0.0018 ± 0.0029	0	0.0143 ± 0.0026	0.002 ± 0.003	0
time of periastron passage, T (HJD)	2456899.0 ± 5.9		2453507.96 ± 0.71	2456985.3 ± 3.7	
longitude of periastron, ω_1 ($^\circ$)	$38. \pm 109$		46 ± 13	181 ± 68	
time of maximum velocity T_0 (HJD)		2456916.571 ± 0.012			2453563.81 ± 0.02
$M_1 \sin^3 i$ (M_\odot)	1.126 ± 0.068	1.126 ± 0.065			
$M_2 \sin^3 i$ (M_\odot)	0.655 ± 0.022	0.655 ± 0.021			
$a_1 \sin i$ (10^6 km)	9.487 ± 0.027	9.488 ± 0.027	9.374 ± 0.024	9.33 ± 0.02	9.33 ± 0.02
$a_2 \sin i$ (10^6 km)	16.31 ± 0.43	16.32 ± 0.42			
$a \sin i$ (R_\odot)	37.09 ± 0.62	37.09 ± 0.60			
mass ratio, $q = M_2/M_1$	0.582 ± 0.016	0.582 ± 0.15			
mass function, $f(M) = (M_2 \sin i)^3 / (M_1 + M_2)^2$			0.0854 ± 0.00066	0.0841 ± 0.0006	0.0841 ± 0.0006
RMS velocity residuals, σ_1 (km s^{-1})	0.29	0.28	0.45	0.30	0.30
RMS velocity residuals, σ_2 (km s^{-1})	4.8	4.6			
light ratio at 519 nm	0.0139	0.0139			

 NOTE. — ^aThese radial velocity data sets are the complete sets and include data contaminated by starspots.

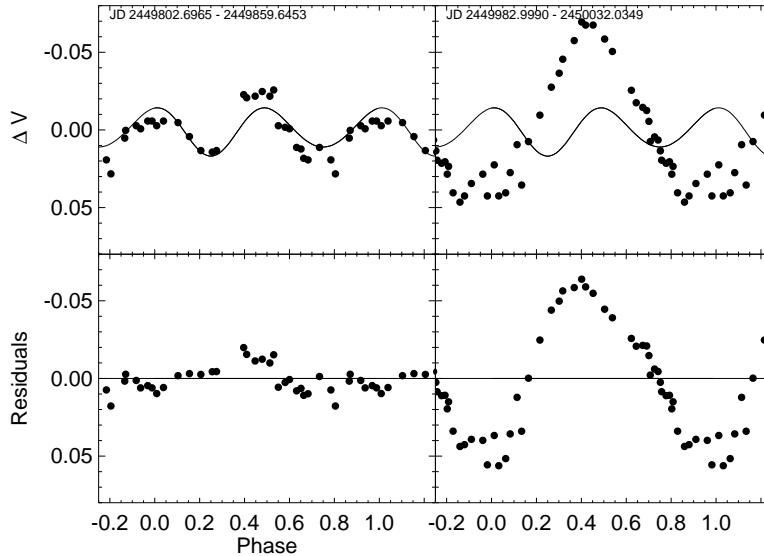


FIG. 6.— Average-subtracted, differential light curves of σ Gem for Johnson V magnitudes plotted for JD 2449802.6965 – 2449859.6453 (left; Segment 8, Set 45 (Kajatkari et al. 2014)) and JD 2449982.9990 – 2450032.0349 (right; Segment 9, Set 1 (Kajatkari et al. 2014)). The top panel contains a plot of the APT data sets (circles) and the model ellipsoidal variations created with ELC for the orbital parameters of σ Gem and best-fit gravitational darkening coefficient $\beta = 0.02$ (solid line). The bottom panel contains the residuals of the APT light curve with the ellipsoidal variation signature removed (circles).

TABLE 8
JOHNSON PHOTOMETRIC ELLIPSOIDAL VARIATIONS MODELS OF σ GEM

Phase	U	B	V	R	I	J	H	K
0.000	-0.018	-0.016	-0.014	-0.013	-0.012	-0.011	-0.011	-0.010
0.011	-0.019	-0.016	-0.014	-0.013	-0.012	-0.011	-0.011	-0.010
0.022	-0.019	-0.016	-0.014	-0.013	-0.012	-0.011	-0.010	-0.010
0.033	-0.018	-0.015	-0.014	-0.012	-0.011	-0.011	-0.010	-0.009
0.044	-0.017	-0.015	-0.013	-0.012	-0.011	-0.010	-0.009	-0.009

NOTE. — Table 8 is published in its entirety in the electronic edition. Orbital phase is based upon our orbital parameters listed in Table 6. Notably, $T_0 = 2453583.98$ (HJD), $P_{\text{orb}} = 19.6027$ days.

signature of two spots on opposite sides of the star (the basis for the active longitudes claims) nearly completely disappears (see Figure 6). The second epoch (“Segment 9, Set 1” of Kajatkari et al. (2014)) is dominated by one spot and the ellipsoidal variations are not discernible. Nonetheless, future starspot modelers should account for the underlying ellipsoidal variations before performing detailed light curve analysis or surface brightness inversions. We have included our calculation of the expected ellipsoidal variations for Johnson $UBVRIJK$ in Table 8 to assist future workers—a full re-analysis of the star spot properties is beyond the scope of this paper.

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