

A Pan-STARRS 1 study of the relationship between wide binarity and planet occurrence in the *Kepler* field

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ABSTRACT

The NASA *Kepler* mission has revolutionised time-domain astronomy and has massively expanded the number of known extrasolar planets. However, the effect of wide multiplicity on exoplanet occurrence has not been tested with this dataset. We present a sample of 401 wide multiple systems containing at least one *Kepler* target star. Our method uses Pan-STARRS 1 and archival data to produce an accurate proper motion catalogue of the *Kepler* field. Combined with Pan-STARRS 1 SED fits and archival proper motions for bright stars, we use a newly developed probabilistic algorithm to identify likely wide binary pairs which are not chance associations. As by-products of this we present stellar SED templates in the Pan-STARRS 1 photometric system and conversions from this system to *Kepler* magnitudes. We find that *Kepler* target stars in our binary sample with separations above 6'' are no more or less likely to be identified as confirmed or candidate planet hosts than a weighted comparison sample of *Kepler* targets of similar brightness and spectral type. Therefore we find no evidence that binaries with projected

separations greater than 3,000 AU affect the occurrence rate of planets with $P < 300$ days around FGK stars.

Key words: binaries: visual, astrometry: proper motions, stars: planetary systems

1 INTRODUCTION

The *Kepler* Mission (Borucki et al. 2010) has to date identified 3697 exoplanet candidates (Rowe et al. 2015 and references therein). The confirmed candidates and inferred false positive rates have been used to determine the fraction of stars with shorter period planets (Howard et al. 2011; Fressin et al. 2013) with particular attention paid to the number of Earth-like planets around M dwarfs (Morton & Swift 2014; Dressing & Charbonneau 2015). The determination of the parameters of transits exoplanets depends heavily on the observed properties of the host star. Hence the study of exoplanet hosts' metallicity, radius and temperature has been a fruitful endeavour for several groups (Mann et al. 2012, 2013b,a; Muirhead et al. 2014; Newton et al. 2015). One key problem with late-type stars is the significant pollution of the dwarf sequence by background giants. While these can be readily identified through spectroscopy (Mann et al. 2012) this is hard to do for large numbers of stars. Because planets are difficult to impossible to find around giant stars using transit searches (and because evolved stars can cause errors in planet property determinations, Bastien et al. 2014; Wang et al. 2015), planet occurrence calculations require an estimate of the number of evolved stars in the survey (Mann et al. 2012). Therefore, calculating the planet fraction requires an accurate characterisation of the tens of thousands of *Kepler* target stars with no detected planet. Mann et al. (2013a) showed that reduced proper motion is an efficient dwarf/giant discriminator for late K and M dwarfs which can be used on much larger samples of stars.

The quality of proper motions in the *Kepler* field varies with the magnitude of the target star. For the brightest objects either Hipparcos (to $V \approx 9$) and Tycho (to $V \approx 11.5$) provide proper motion measurements with uncertainties of the order of a few millarcseconds per year. Astrometric observations in the UCAC 4 catalogue provide proper motion errors below 10 millarcseconds per year for stars down to $R \approx 16$. This covers the majority of the *Kepler* targets but will miss any faint common proper motion companions to these *Kepler*

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targets. Digitised archival photometric surveys such as SuperCOSMOS (Hambly et al. 2001) and USNO-B (Monet et al. 2003) provide proper motions for these fainter stars. While these surveys have the advantage of long time baselines, their photographic nature means that each individual epoch is of poorer astrometric quality than modern CCD surveys. Additionally the most extensive proper motion survey based on photographic plates, the Superblink survey (with errors of approximately 8 milliarcseconds per year) only has a few subsets of data which are publicly available: high proper motion stars in the Northern hemisphere (Lépine & Shara 2005) and bright M dwarfs (Lépine & Gaidos 2011).

Wide stellar companions are common in the field, with >25% of nearby solar type stars having companions wider than 100 AU (Raghavan et al. 2010). Among young stars, disk frequency (Kraus et al. 2012) and disk mass (Pascucci et al. 2007; Harris et al. 2012) are suppressed for close binaries but the components of systems wider than 100 AU have similar disk properties to single stars. This suggests that wide binaries play a minor role in the initial sculpting of planetary systems. Raghavan et al. (2006) found 24 companions wider than 100 AU amongst a sample of 131 radial velocity exoplanet hosts. This equates to a wide companion fraction of $18.3 \pm 3.7\%$, slightly lower than the $\sim 25\%$ of solar-type stars in Raghavan et al. (2010) which have companions wider than 100 AU. Finally, Wang et al. (2014) studied binary systems closer than 1500 AU finding a significant deficit of planets around the closest binaries with the effect dropping significantly becoming very small for systems wider than 100 AU.

Recent work by Kaib et al. (2013) suggests a different mechanism by which binaries with semi major axes of 1,000 AU or so can affect planetary systems. While orbiting through the Milky Way's gravitational potential, the orbital parameters of wide binary companions can be modified by gravitational interactions with passing stars driving them to more elliptical orbits and sending the companions close to their primary at periastron. Such close passage can disrupt the outer planets in such systems and induce planetary migration and/or ejection. The magnitude of this effect can be tested by comparing the planet occurrence rate around wide binary components with that around field stars. However this is complicated by coincident alignments of unrelated pairs of stars often mimicking true wide binaries. Accurate kinematic measurements and statistical models are required to separate true wide binaries from coincident pairings.

Wide binaries are also useful astrophysical probes, ideal for calibrating and testing empirical stellar property estimators. Gyrochronology, the process of estimating a star's age

from its rotation period, is one of the most valuable techniques in stellar astronomy (Barnes 2007). This can be used to derive ages for exoplanet host stars (e.g. Walkowicz & Basri 2013) or ultracool companions to intermediate mass stars (e.g. Dupuy et al. 2009). These relations are often calibrated using open clusters of stars as anchor points. Each cluster provides a few hundred stars of different types, all of the same age. These discrete anchor points have the disadvantage that there are relatively few well studied, older open clusters with well determined rotation periods for their stellar members. Meibom et al. (2015) recently provided rotation periods for the 2.5 Gyr-old cluster NGC 6819. More field-age tests of gyrochronology relations will improve stellar age estimates even further, but each cluster still only represents a single age. Wide binary systems, however, cover a wide range of ages. While their age is not known independently as with clusters, they can still be used as a powerful test of gyrochronology at ages outside the range of known clusters. Wide binary systems are coeval systems and thus their components should have similar ages measured from gyrochronology. *Kepler* provides excellent photometry for rotation period calculation so any wide binaries of two *Kepler* target stars can be used to test gyrochronology relations. Conversely, gyrochronology relations can be used to test if wide binary systems are true physical pairs.

In this paper we present a proper motion survey of the *Kepler* field. We combine archival datasets from wide-field public surveys and target UKIRT observations with new Pan-STARRS1 astrometry to produce accurate proper motions for *Kepler* target stars. This enables us to combine the excellent astrometric accuracy of a modern CCD-based survey with the long time baseline of archival plate data. We then use our Pan-STARRS1 plus archival data proper motion catalogue, along with SED fits using Pan-STARRS1 and 2MASS photometry, to select a population of wide binary pairs where at least one component is a *Kepler* target star. We then use these binaries to test whether wide binarity has an effect on exoplanet occurrence.

2 DATASETS

2.1 Pan-STARRS 1 data

A 1.8 m high altitude survey telescope, Pan-STARRS 1 (Kaiser et al. 2002) recently completed its full three and a half year survey operations on Haleakala on Maui in the Hawaiian Islands. This consisted of a suite of surveys of different cadences and depths. The *Kepler*

field was included in the 3π survey which covers the full sky north of $\delta = -30^\circ$ in five filters (g_{P1} , r_{P1} , i_{P1} , z_{P1} and y_{P1} ; Tonry et al. 2012). Each filter has approximately six pairs of observations each separated by roughly half an hour over the course of the survey for any one point on the sky, therefore a typical object will have 30 pairs of observations. The data from Pan-STARRS1 are astrometrically and photometrically reduced and calibrated using the processes outlined in Magnier (2006), Magnier (2007), Schlafly et al. (2012) and Magnier et al. (2013).

2.1.1 Astrometric accuracy

There are three separate uncertainties in our Pan-STARRS1 astrometric uncertainty model: centroiding errors caused by photon noise, the Pan-STARRS1 internal systematic floor and the systematic floor caused by comparing Pan-STARRS1 with other surveys.

The astrometric centroiding error typically takes the form $\sigma_{pos} = \sqrt{a^2 + b^2\sigma_{mag}^2}$ where σ_{pos} is our positional error in arcseconds, σ_{mag} is the photometric error in magnitudes, a is our systematic floor and b is some conversion factor which depends on the PSF shape and whether the object lies in the source or background dominated noise regime (King 1983). The Pan-STARRS1 PSF takes a form which includes a free parameter (Magnier et al., in prep.). Thus we cannot know b without knowing the exact PSF shape for every object. As this would be cumbersome and computationally impractical to implement, we take an empirical approach. We extracted all r_{P1} detections in a 1×1 degree patch centred on $R.A. = 18^h 42^m$ $Dec. = 46^\circ 30^m$ and measured the scatter in each object's position about the mean position. We then plotted this as a function of the mean photometric error for the detections of that object (see Figure 1). We found that a model of the form $\sigma_{pos} = \sqrt{0.015^2 + \sigma_{mag}^2}$ was a good fit to the data. Hence we adopt 15 milliarcseconds as our Pan-STARRS1 internal systematic floor and set our photon-noise positional errors to be the same as the magnitude error (i.e. $b = 1$).

To estimate the systematic error added by comparing Pan-STARRS1 data with other surveys we again took another 1×1 degree test area (this time centred on $R.A. = 19^h 14^m$ $Dec. = 38^\circ 30^m$ as our first field has no SDSS data) and examined the positional scatter between Pan-STARRS1 and three external surveys (SDSS, Aihara et al. 2011; 2MASS, Skrutskie et al. 2006 and P.I. UKIRT data of the *Kepler* field processed using the same pipeline as UKIDSS data, Lawrence et al. 2007). Selecting only bright objects which lay

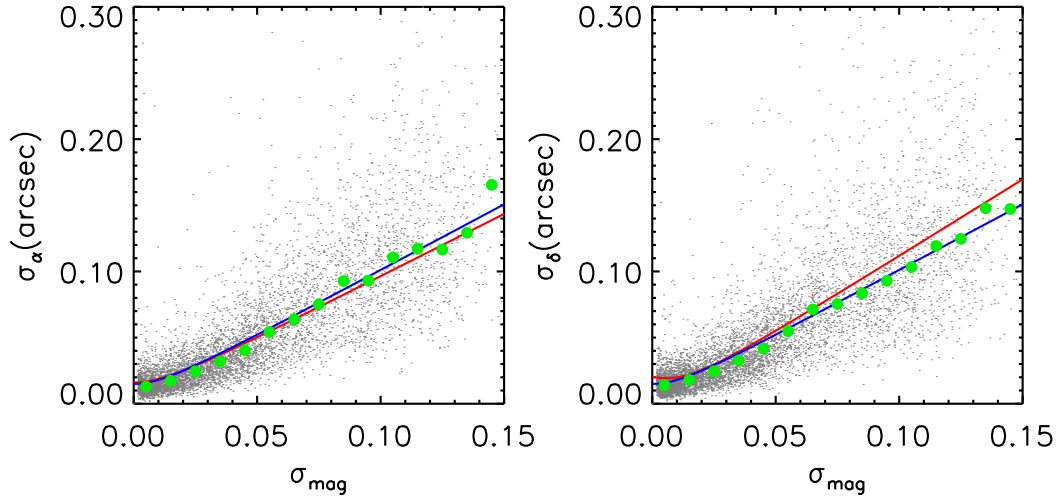


Figure 1. Astrometric scatter of r_{P1} detections about mean object positions as a function of mean photometric error for the object's r_{P1} detections. The red line shows a least squares fit of the form $\sigma_{pos} = \sqrt{a^2 + b^2 \sigma_{mag}^2}$ and the blue line shows our assumed astrometric model with $a=15\text{mas}$ and $b=1$. Green dots are the median scatter in each 0.01 mag bin.

in the magnitude range of each of the surveys where the astrometric uncertainty is dominated by systematic errors ($15 < r < 16$ for SDSS and Pan-STARRS 1, $11 < J < 15$ for 2MASS and UKIRT) we examined what the scatter about zero was for each pair of surveys. This provided us with six linear combinations of the systematic variances of the surveys. We solved these as a set of equations by singular value decomposition finding that the systematic variance of Pan-STARRS 1 data compared to external datasets is approximately 38 milliarcseconds (including our internal systematic floor of 15 milliarcseconds). Once the internal systematic floor is subtracted in quadrature, we find a systematic error going from the Pan-STARRS 1 reference frame to some external reference of 35 milliarcseconds. Both external floor and internal floor are used later in our calculation of proper motions.

2.2 UKIRT data

We made use of observations which were made of the *Kepler* field using WFCAM (Casali et al. 2007) on the UK Infrared telescope which are now public. These data came from program U/09A/2 (PI Lucas) observed on July 11th–13th 2009. Observations were done in the J_{MKO} -band in a similar set-up to the UKIDSS Large Area Survey J_{MKO} observations (2×2 microstepping, one telescope offset, 40 s total integration, Lawrence et al. 2007) with reduced data catalogues available as FITS tables in the WFCAM Science Archive (Hambly et al. 2008) having been processed by the WFCAM pipeline (Irwin et al. 2004).

2.3 Public survey data

We included astrometry from multiple public surveys in our work. We drew data from the Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006), the Sloan Digital Sky Survey (SDSS DR8; Aihara et al. 2011), the USNO-B digitisation of photographic plate data (Monet et al. 2003) and the WISE All-Sky Survey (Wright et al. 2010). The astrometry and photometry from these surveys were all extracted from the *Vizier* online archive service with the exception of WISE dataset which was downloaded from the NASA/IPAC Infrared Science Archive. To avoid spurious faint photographic plate detections affecting our proper motion fits we ignored any detections in USNO-B fainter than $B_J = 19$, $R = 19$ and $I_N = 18$. This last cut is similar to that used in Deacon et al. (2005) and Deacon & Hambly (2007).

For archival surveys we used the astrometric error estimates provided by the relevant survey archives with the exception of SDSS. For this survey we used the error model used by Kraus & Hillenbrand (2007) which features a 40 milliarcsecond systematic floor (similar to that found by Pier et al. 2002) and photon noise term that scales as the quoted magnitude uncertainty.

3 CALCULATION OF PROPER MOTIONS

We divided the *Kepler* field into 1×1 degrees chunks and in each of these areas extracted Pan-STARRS 1 data from the second large-scale reprocessing of the data (PV2) using scripts written in the *Desktop Virtual Observatory* shell language (Magnier et al. 2008). We extracted as our base catalogue the average catalog objects in the target area. This average object catalogue consists of photometric and astrometric properties derived from all observations of each individual object and includes initial Pan-STARRS 1-only proper motions. As some objects will have proper motions that will take their archival detections outside our pairing radius of $1''$, we used the calculated Pan-STARRS 1-only proper motions to find the likely position of each object at the epochs of each archival non-Pan-STARRS observation. We then searched around these positions with a pairing radius of $1''$ to identify the objects' detection in each survey. This procedure is vulnerable to spurious proper motions in the Pan-STARRS 1 database so we do not applying the proper motion “rewind” procedure for any object with a Pan-STARRS 1 proper motion $\chi_v^2 > 10$.

To calculate our proper motions we constructed a covariance matrix taking into account that the Pan-STARRS 1 measurements will have three errors associated with them, photon

noise, the systematic floor of relative Pan-STARRS1 astrometry and the systematic error between Pan-STARRS1 and other surveys. The Pan-STARRS1 measurements will have errors correlated with each other with a covariance of the square of the Pan-STARRS1 to other surveys systematic floor. No other survey data was given off-diagonal terms in the covariance matrix. We then calculated the proper motion for each object based on the measurements and our covariance matrix.

We performed our fit in three stages. Having examined a number of preliminary fits we found that these were often offset by one or two Pan-STARRS data points which were significant outliers. Performing a 3σ clip around our initial fit would be the most obvious solution. However this initial fit would be strongly affected by the outlier points and hence may exclude many valid datapoints which were outliers to this erroneous initial proper motion fit. Hence we performed an initial clipping removing any Pan-STARRS1 data point which was more than a 10σ outlier (i.e. a 10σ outlier in the quadrature sum of the number of standard deviations the object was an outlier from the fit in R.A. and Dec.). After this, we refitted our data and applied a more stringent 3σ cut again only on the Pan-STARRS1 data. This was followed by a final refitting of our proper motions. Figure 2 shows an example of the clipping and fitting process on one *Kepler* target.

3.1 Selecting the proper motions of Kepler targets

Often there are multiple potential Pan-STARRS1 matches for a particular Kepler target. We followed the following procedure to select the most appropriate matches. Firstly we selected only objects which had g_{P1} , r_{P1} or i_{P1} magnitudes within 2 magnitudes of the g , r and i magnitudes of the Kepler target in the Kepler Input Catalog. This excludes matches with obviously spurious faint sources. We then matched each Kepler target with the predicted Epoch=1999.0 positions of all objects within $3''$ and implemented the following procedure for multiple matches.

(i) We asked how many of those matches are within $1''$ and have r magnitudes in the KIC that agree with Pan-STARRS1 r_{P1} to within 0.2 mag.

(a) If there is one such match we use this one

(b) If there are multiple of these we use the one with the most measurements in its astrometric fit

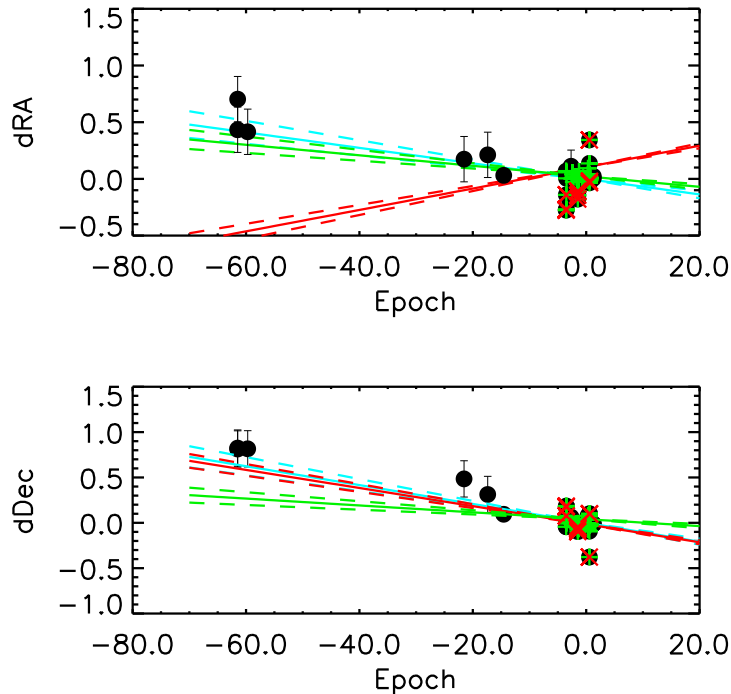


Figure 2. An example of the fitting and clipping process for the star KIC 6878408. The initial fit is the red line, the second fit after the 10σ clip is the green line and the final fit after the 3σ clip is the blue line. The dashed lines represent the 1σ confidence limits on the proper motion in each fit. Note points crossed out with red are excluded by the 10σ clip and those crossed out in green are excluded by the 3σ clip.

(ii) If not we ask how many of the matches have r magnitudes in the KIC that agree with Pan-STARRS1 r_{P1} to within 0.2mags

- (a) If there is one such match we use this one
- (b) If there are multiple of these we use the closest one

(iii) Otherwise we use the closest positional match

3.2 Proper motion results

Proper motions are presented in Table 1. In our remaining analysis we define significant proper motions as being more significant than 5σ and good quality proper motions as being calculated from 10 or more measurements, having a USNO object within $6''$ and one PS1 object within $3''$, being fainter than $r_{P1} = 14.5$ mag. (i.e. not saturated), brighter than $r_{P1} = 19$ mag and having a reduced χ^2 less than 4. This results in proper motions which are not confused by crowding, saturation or outliers and likely involving a time baseline of over 50 years. Figure 3 shows the distribution of all objects in our sample with unsaturated photometry (left) and photometry for unsaturated *Kepler* targets (right) in a $r_{P1} - i_{P1}$

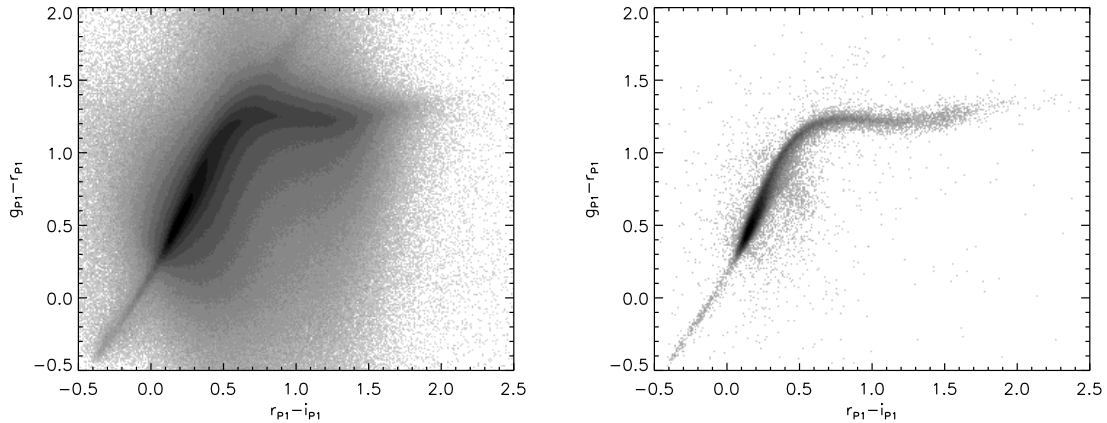


Figure 3. The \log_{10} density distribution of all unsaturated ($r_{P1} > 14.5$ mag.) sources (left) and unsaturated *Kepler* targets (right) on a colour-colour plot. Note the concentration of *Kepler* targets around (0.2,0.5) in the *Kepler* targets plot. This is due to the large number of F and G stars in the *Kepler* target list. Note also the lack of an extension of the F, G, K locus after the stellar locus takes a turn at (0.5,1.2) with the onset of TiO absorption in M dwarfs. This suggests few *Kepler* targets in our sample are significantly reddened.

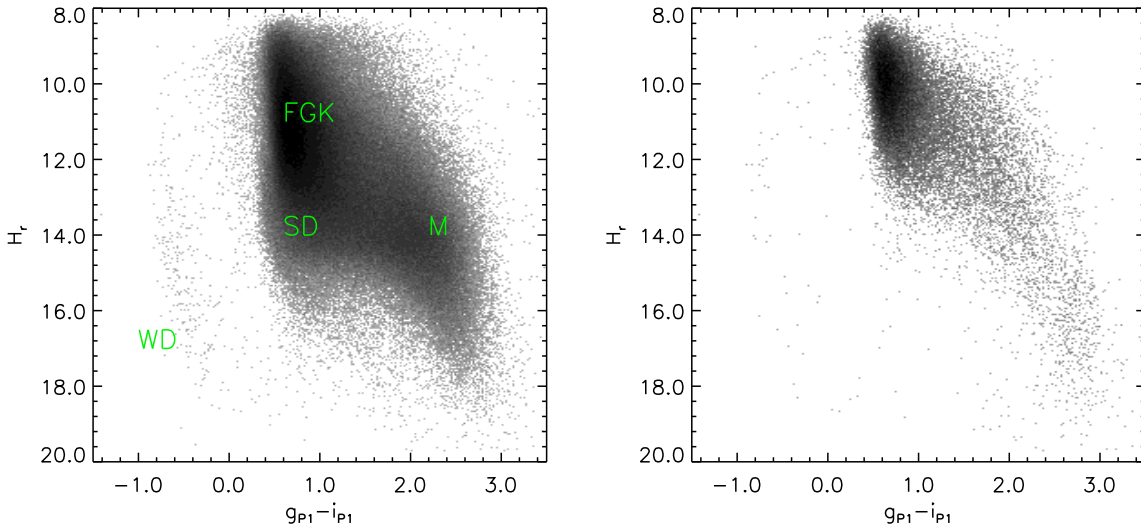


Figure 4. The \log_{10} density distribution of reduced proper motion vs. colour for all sources with good quality 5σ proper motions (left) and *Kepler* targets meeting the same criteria. Note the clear white dwarf (marked WD), F, G, K dwarf (marked FGK), subdwarf (marked SD) and M dwarf (marked M) loci on the left-hand diagram. The right-hand plot is dominated by F, G and K stars, showing the known preference in *Kepler* target selection.

vs. $g_{P1} - r_{P1}$. This shows that there are relatively few *Kepler* targets which appear to be reddened beyond the stellar locus. Figure 4 shows reduced proper motion diagrams ($H_r = r_{P1} + 5 + 5 \log_{10} \mu$) for both all objects with good quality significant proper motions (left) and *Kepler* targets passing the same criteria (right). Clear F, G, K dwarf, M dwarf, white dwarf and subdwarf loci can be seen in the left-hand diagram, but the *Kepler* targets are disproportionately dominated by F, G, and K dwarfs. This is unsurprising as these were the primary target class in *Kepler*'s search for Earth-like planets (Batalha et al. 2010).

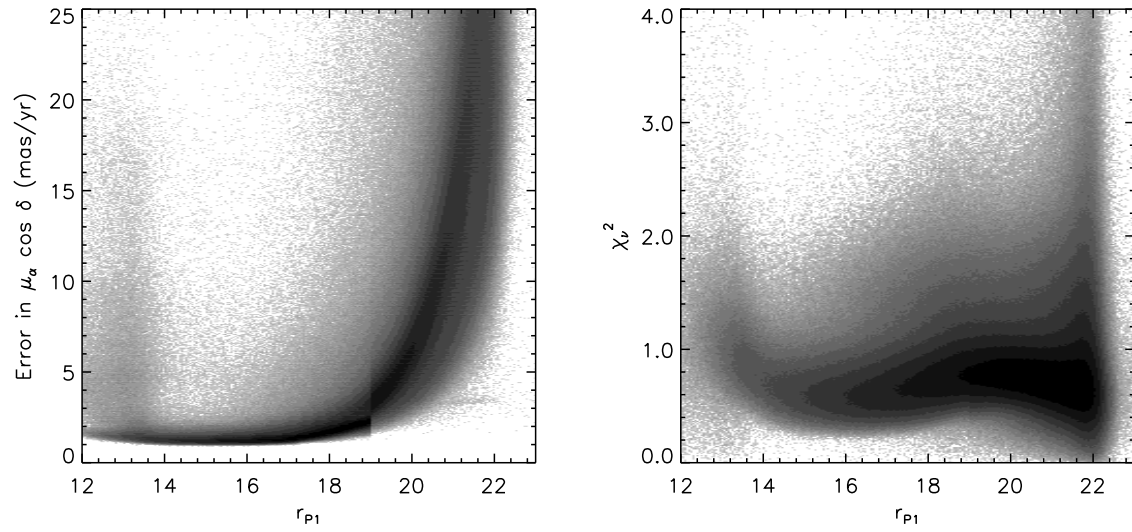


Figure 5. The \log_{10} density distribution of proper motion error vs. magnitude (left) and reduced χ^2 vs. magnitude (right) for all sources. Note the discontinuity in proper motion error at $r_{P1} = 19$. This is because we excluded matches with USNO detections at this magnitude to prevent matches with faint, spurious USNO detections. Note also both distributions become warped brighter than $r_{P1} = 14$ due to saturation. Our reduced χ^2 plot peaks at a value slightly below 1, indicating our astrometric errors are slightly over-estimated.

Table 1. Proper motions for unsaturated *Kepler* targets along with their Pan-STARRS1 magnitudes. The penultimate column flags if this object passes the "good" proper motion test previously set out in the text and the final column shows *Kepler* Objects of Interest which have passed or failed our vial proper motion inspection. The number of objects within 3 arcseconds and 6 arcseconds in both Pan-STARRS1 and USNO are also quoted. Objects marked with † have multiple PS1 matches within 3'' but the astrometric solution selected is for an object within one arcsecond with the same magnitude as the Kepler target and is thus believed to be reliable. The full table is available electronically.

KID	R.A. (Eq=J2000 Ep=2013.0)	Dec.	$\mu_{ra} \cos \delta$ (mas/yr)	μ_{δ} (mas/yr)	n_{meas} (mas/yr)	g_{P1} (mag/yr)	r_{P1}	i_{P1} (mag.)	z_{P1} (mag.)	y_{P1} (mag.)	χ^2_{ν} (mag.)	$n3_{PS1}$ (mag.)	$n6_{PS1}$	$n3_{USNO}$	$n6_{USNO}$	flag	v
757450	19 24 33.03	+36 34 38.6	12.08±1.65	2.85±1.54	59	15.773±0.001	15.087±0.001	14.844±0.001	14.737±0.001	14.679±0.001	0.485	1	1	1	1	1	1
891916	19 23 49.98	+36 41 11.8	8.79±1.61	2.26±1.5	57	15.179±0.001	14.691±0.001	14.512±0.001	14.439±0.001	14.392±0.001	0.626	1	2	1	1	1	1
892718	19 24 34.01	+36 38 53.9	-4.83±1.39	-4.04±1.27	65	16.515±0.002	15.842±0.001	15.586±0.001	15.451±0.001	15.373±0.002	0.553	1	1	1	2	0	0
892772	19 24 36.80	+36 40 43.8	-8.16±1.3	0.95±1.19	74	15.77±0.001	15.05±0.001	14.772±0.001	14.627±0.001	14.525±0.001	0.72	1	2	1	1	1	1
892832	19 24 39.15	+36 40 27.5	-13.11±1.32	-12.47±1.2	81	16.304±0.002	15.624±0.001	15.367±0.001	15.252±0.001	15.177±0.001	0.326	1	1	1	2	0	0
892834	19 24 39.20	+36 37 39.3	2.38±1.39	2.28±1.28	69	15.919±0.001	15.134±0.001	14.863±0.001	14.731±0.001	14.657±0.001	0.617	1	2	1	1	1	1
892882	19 24 41.58	+36 41 54.1	1.89±2.47	8.43±2.23	23	15.53±0.001	14.862±0.001	14.638±0.001	14.542±0.001	14.483±0.001	1.363	1	3	1	1	1	1
892911	19 24 43.13	+36 40 14.1	10.57±1.93	0.18±1.86	18	16.081±0.001	15.601±0.001	15.433±0.001	15.366±0.001	15.314±0.002	1.046	2	2	1	1	0	0
892946	19 24 45.03	+36 40 21.0	8.44±1.85	3.73±1.76	25	16.237±0.001	15.804±0.001	15.663±0.001	15.622±0.002	15.586±0.002	1.285	1	2	2	2	0	0
893033	19 24 49.82	+36 40 03.8	19.18±1.32	-8.03±1.21	77	15.976±0.001	15.197±0.001	14.899±0.002	14.776±0.001	14.699±0.001	0.442	1	1	1	2	0	0

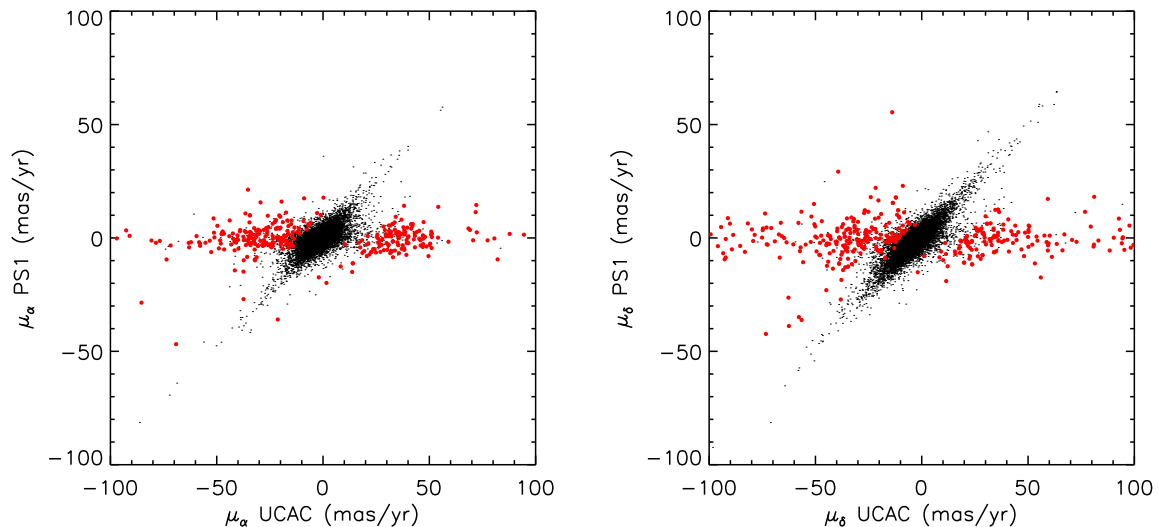


Figure 6. Comparison of our Pan-STARRS1 proper motions with those from the UCAC4 catalogue (Zacharias et al. 2013). Only points which pass our proper motion quality cuts are included. Red points show 5σ disagreements between our proper motions and UCAC4 in R.A (left panel) or Dec. (right panel). The majority of data points fall along a one-to-one relation between the two proper motions. A number of objects have low proper motions in our calculations but high UCAC proper motions

3.3 Verification of proper motions

To verify our proper motions we compared our selection of reliable proper motions with the values from the UCAC4 survey (Zacharias et al. 2013; see Figure 6). This dataset does not use Schmidt survey plates so it represents an independent dataset. We found that our values generally compared well with the UCAC values. However, there was a small number (3.3%) of objects for which we measure much lower proper motions compared to UCAC. A visual inspection of a subset of these objects suggests that these objects do not have significant proper motions and often have another star at a $6\text{--}10''$ distance suggesting that these high proper motions are an artefact of object confusion in UCAC. As an additional check, we individually inspected the astrometric solutions for 3800 Kepler Objects of Interest listed in Rowe et al. (2015) for which we had a proper motion solution. We then made a by-eye judgement on how reliable the proper motion solution was compared to the distribution of astrometric points. We found that 94% of our proper motions passed this test. Of the objects which failed visual inspection, approximately three quarters had proper motions which would not meet our reliability cuts (high reduced χ^2 , few measurements, etc.). This KOI visible checking is listed as a parameter in Table 1.

3.3.1 *White dwarf Kepler Targets*

Proper motion catalogues with high levels of contamination due to spurious proper motions often present reduced proper motion diagrams with white dwarf loci which are heavily contaminated by FGK stars with erroneously high proper motion measurements. Hence the population of the white dwarf locus is a useful test of the reliability of a proper motion catalogue. To identify a sample of *Kepler* target white dwarfs we selected unsaturated objects which had good, significant proper motions, were redder than $g_{P1} - i_{P1} = -2$ (to exclude objects with erroneously blue photometry), met the condition $H_r > 10 \times (g_{P1} - i_{P1}) + 12$ and whose proper motion solutions passes a visual inspection. These objects are shown in Table 2. Most of the intrinsically brighter objects passing this cut (those with low reduced proper motion) are sdO or sdB stars based on literature classifications (see Table 2). This cut also misses intrinsically bright white dwarfs which will be distant enough that they would not pass our proper motion significance cut. Twelve of our fifty candidate white dwarfs have been previously identified as white dwarfs, dwarf novae or cataclysmic variables. We do not see significant contamination from objects spectrally classified as FGK stars.

4 IDENTIFICATION OF LIKELY BINARY COMPANIONS

4.1 Probabilistic binary selection

Stars in wide binary systems are typically found due to their common proper motion. Identification of such pairings is complicated by the presence of nearby stars, which by chance, have the same proper motion as each other despite having no physical connection. These coincident pairings are often removed with hard cuts or using by-eye judgements. Two better defined processes for binary star selection come from the work of Lépine & Bongiorno (2007) and Dhital et al. (2010). The former empirically defines a region of proper motion and separation space where coincident pairings are unlikely. The latter method uses a model of the Galaxy to define the probability that an unrelated star would have a particular proper motion and position difference from the target primary star. Here we define a hybrid of these two systems which calculates the probability that an object is a true binary based on an estimate of the background star population, but unlike Dhital (2010), our coincident pairing estimates are purely empirical. However unlike the two previously mentioned methods, we take the likely distribution of the wide binary population into account as well. Similar methods of separating objects from a background population have been used to

Table 2. ¹Ostensen et al. (2010), ²Ostensen et al. (2011b), ³Reed et al. (2011), ⁴Girven et al. (2011), ⁵McNamara et al. (2012), ⁶Hoffmeister (1966), ⁷Ostensen et al. (2011a), ⁸Howell et al. (2013), ⁹Bischoff-Kim & Østensen (2011), ¹⁰Scaringi et al. (2013), ¹¹Feldmeier et al. (2011), ¹²Baran et al. (2011), ¹³Gianninas et al. (2011), ¹⁴Hog et al. (2000), ¹⁴Greiss et al. (2014).

<i>Kepler_ID</i>	<i>Other name</i>	<i>Literature classification</i>	<i>Reference</i>
3353239	2MASS J19364633+3825268	sdB	2
3527751	2MASS J19033701+3836126	sdB	3
3561157			
3629120	SDSS J190410.70+384518.1	WD cand	4
3663415			
3751235			
4357037	SDSS J191719.16+392718.8	WD cand	4
5077438	TYC 3123-1545-1		14
5544772			5
5951261			5
6042560			
6212123			
6672883			
7346018			
7594781			
7659570	V* V344 Lyr	DN	6
7797992			
8075923			
8077281	2MASS J18501686+4358284	B	1
8210423			
8244398			
8395780			
8420780	GALEX J192904.6+444708	DB	7
8490027		CV cand	8
8626021	WD J1929+4447	DB	9
9071514		CV	10
9139775		DA ¹	
9228724			
9391127			
9535405		DA	11
9569458		sdB	1
9818160			
10066680			
10081214			
10149875			
10290697			
10462707		sdB	2
10590313		LSPM J1910+4750	
11176123			
11191648			
11351218			5
11357853		sdOB	1
11402999			
11509531			
11558725	2MASS J19263411+4930296	sdB	12
11604781	2MASS J19140898+4936410	DA	2
11822535	WD 1942+499	DA1.4	13
11911480		DA/ZZ Ceti	14
12021724	2MASS J19441275+5029393	sdB	2
12353867			

identify members of open clusters using their proper motion such as Deacon & Hambly (2004), Kraus & Hillenbrand (2007) and Sanders (1971). While these methods fit distribution functions to the background population, we use the Lépine & Bongiorno (2007) method of offsetting the position of the target primary and then calculating the density of objects

with similar proper motions, positions, distances and masses to the candidate secondary. Appendix A outlines the mathematics used in our probability calculations.

4.2 Testing the binary probability algorithm

To test that our method for binary association probability is accurate we used a sample of bright stars from our *Kepler* proper motion catalogue. These had proper motions from UCAC4 (Zacharias et al. 2013) or PPMXL (Roeser et al. 2010) and photometric distances and mass estimates calculated using the methods of Kraus & Hillenbrand (2007) and Kraus et al. (2014b). We assumed a log-flat separation distribution (Opik 1924)¹, a flat mass ratio distribution (Raghavan et al. 2010) and set the binary fraction wider than 100 AU to be 25%. This latter number is similar to the number of binaries wider than this limit found around nearby solar-type stars by Raghavan et al. (2010). We did not set a lower proper motion limit and considered only pairs with observed separations below five arcminutes and projected separations below 10,000 AU. We used this technique to calculate companionship probabilities for these pairs and then flagged a subsample of these pairs for follow-up observations.

A total of 26 wide binary candidate systems (52 stars) were observed with the Tull Coude spectrograph (Tull et al. 1995) at the Harlan J. Smith 2.7m Telescope at McDonald Observatory on the nights of August 8, 9, 10, and 11, 2014. Exposure times varied from 120s to 1200s, which was sufficient to achieve a median SNR (per resolving element) of > 30 . The resulting spectra cover 3800Å to 10500Å, with a gap from 4550Å to 4750Å due to internal reflection, and at a resolution of $R \simeq 60,000$. Pair components were always observed within minutes of each other to reduce systematic sources of error (e.g., atmospheric changes, telescope flexure, Earth's motion). Data were reduced using standard IRAF reduction tools, including bias and flat field correction, extraction of the 1 dimensional spectrum, and wavelength calibration. An additional wavelength correction was applied by cross correlating the observed telluric lines to a model atmosphere as described in Gullikson et al. (2014). We cross-correlated the observations of our target pair components with each other to determine the radial velocity difference. These values for our observed pairs (along with their binarity probability p_{bin}) are listed in Table 3. The errors in radial velocity difference come

¹ This is flatter than the typical log-normal separation distribution used for binaries closer than 100 AU but is appropriate as the binaries we identify in this work are typical a few thousand AU in separation

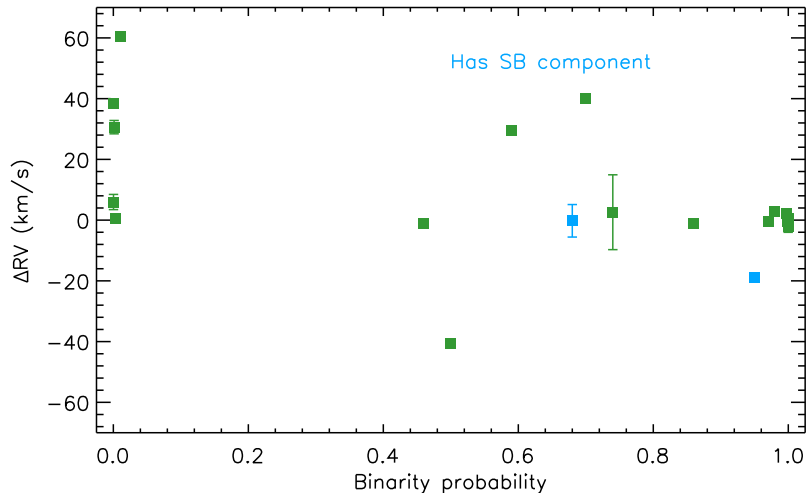


Figure 7. The observed radial velocity differences for a series of candidate binaries with a range of calculated binary probabilities. These objects had proper motions from UCAC4 and SED fits from the catalogue built for Kraus et al. (2014b). Two of our pairs have a component which is a spectroscopic binary. This will induce offsets in their radial velocity. Note also the large fraction of objects which have high binarity probabilities and common radial velocities.

from the cross correlation (with a 4σ clipping) but do not include other potential sources of random and systematic error (e.g, additional undetected companions) and hence they are likely to be an underestimate of the error of the radial velocity difference. Figure 7 shows our results with two pairs flagged as containing a spectroscopic binary (i.e. are hierarchical triples). These were KIC 9085834 and KIC 11017620, which both produced a double-peaked correlation function or showed pairs of stellar lines. Nine of the pairs we classify as having $p_{bin} > 0.6$ have radial velocity differences less than 4 km/s. One of these contains spectroscopic binary while another three objects in the same probability range have radial velocities which are discrepant (excluding spectroscopic binaries). Of the five low probability objects we observed all but one have discrepant radial velocities. This suggests that our method can reliably select true co-moving systems.

4.3 Application to Pan-STARRS 1 *Kepler* targets

The testing of our binary selection algorithm outlined in Section 4.2 only applied to objects with astrometry in UCAC4 and used only photographic plate and 2MASS photometry. We can widen our binary search by looking at fainter companions which have reliable Pan-STARRS1 photometry. To do this we need a series of Spectral Energy Distribution (SED) templates using the process outlined in Appendix B. A by-product of this search allowed us to estimate accurate transformations from Pan-STARRS1 colours to K_p magnitudes.

Table 3. Wide binary candidates with different binary probabilities p_{bin} which had their RV differences measured. Key: *Spectroscopic binary (SB2), †Rapid rotator, Proper motions from ¹Zacharias et al. (2013), ²Roeser et al. (2010).

KIC	$\mu_\alpha \cos \delta$ (mas/yr)	μ_δ (mas/yr)	SpT (subtypes)	Distance modulus (mag)	r (arcsec)	ΔRV (km/s)	p_{bin}
2696938	5.2 ± 2.0^2	-18.3 ± 2.0	F6.7 \pm 1.5	6.31 ± 0.22	5.4	-0.55 ± 0.85	1.00
2696944	1.0 ± 1.8^1	-20.6 ± 2.4	F4.9 \pm 1.4	6.53 ± 0.22			
2992956	16.5 ± 1.4^2	50.2 ± 1.4	G6.5 \pm 1.45	5.13 ± 0.2	16.9	0.47 ± 0.07	1.00
2992960	14.5 ± 1.4^2	51.8 ± 1.4	G6.9 \pm 1.45	5.3 ± 0.19			
4243796	-1.2 ± 0.7^1	-18.7 ± 1.1	K4.1 \pm 0.6	3.91 ± 0.07	154.3	0.52 ± 0.05	0.00
4346953	-9.5 ± 2.0^1	-17.7 ± 2.5	K7.2 \pm 0.15	4.19 ± 0.07			
5790787	4.5 ± 0.9^1	10.6 ± 0.8	F6.9 \pm 1.45	7.45 ± 0.2	27.6	40.07 ± 1.21	0.70
5790807	5.1 ± 0.6^1	9.9 ± 1.0	F3.6 \pm 1.35	6.52 ± 0.23			
6934317	-6.3 ± 0.8^1	-15.5 ± 0.9	G9.9 \pm 1.2	6.46 ± 0.12	44.7	38.29 ± 0.11	0.00
7019341	-11.7 ± 0.6^1	-13.1 ± 0.6	G2.5 \pm 0.75	6.66 ± 0.05			
7090649	16.7 ± 1.8^2	13.6 ± 1.8	G7.0 \pm 2.6	5.97 ± 0.28	9.3	-2.11 ± 1.91	1.00
7090654	17.5 ± 1.2^1	13.3 ± 1.2	F7.2 \pm 1.85	5.80 ± 0.23			
7748234	-6.2 ± 1.6^2	3.9 ± 1.6	F8.2 \pm 0.2	7.14 ± 0.03	38.0	2.6 ± 12.32	0.74
7748238	-7.6 ± 0.9^1	0.7 ± 0.7	F0.0 \pm 1.65	6.8 ± 0.17			
8123664	-5.0 ± 1.5^2	-6.4 ± 1.5	F1.7 \pm 5.15	5.96 ± 2.5	5.6	8.63 ± 0.62	0.00
8123668†	-6.6 ± 1.5^2	-16.9 ± 1.5	F7.5 \pm 4.75	7.70 ± 0.49			
8619322	-4.7 ± 0.9^1	-4.8 ± 1.0	K1.8 \pm 0.85	3.62 ± 0.11	110.8	-40.63 ± 0.07	0.51
8683779	-2.5 ± 1.3^1	-6.3 ± 0.9	K4.4 \pm 0.6	3.87 ± 0.06			
9085833†	-3.8 ± 0.6^1	-15.9 ± 0.8	F4.8 \pm 1.35	7.53 ± 0.21	20.1	-0.22 ± 5.35	0.68
9085834*	-4.3 ± 1.6^1	-13.3 ± 3.4	G2.3 \pm 3.6	8.36 ± 0.32			
9579191	-13.1 ± 2.1^1	10.2 ± 2.0	G5.0 \pm 1.55	7.05 ± 0.15	20.3	-1.27 ± 0.32	0.086
9579208	-11.2 ± 0.8^1	8.8 ± 1.0	F6.4 \pm 1.45	6.39 ± 0.21			
9595822	-2.3 ± 1.0^1	-5.3 ± 0.9	K3.5 \pm 0.25	4.67 ± 0.03	69.0	29.64 ± 0.11	0.59
9534041	-3.5 ± 1.1^1	-6.9 ± 1.0	K1.3 \pm 0.9	4.92 ± 0.12			
9655101	-4.4 ± 0.6^1	-9.3 ± 0.9	K2.2 \pm 0.65	4.73 ± 0.08	65.0	2.7 ± 0.1	0.98
9655167	-4.7 ± 0.7^1	-8.7 ± 1.1	K3.0 \pm 0.45	4.68 ± 0.05			
9777293	-9.8 ± 0.5^1	11.5 ± 0.7	K1.6 \pm 0.85	4.77 ± 0.11	48.9	2.13 ± 0.08	1.00
9777355	-9.4 ± 1.3^1	11.6 ± 0.8	K2.0 \pm 0.7	4.79 ± 0.09			
9912680	11.6 ± 2.0^1	1.0 ± 0.6	G0.4 \pm 2.7	6.53 ± 0.21	7.1	-0.5 ± 0.27	0.97
9912690	14.5 ± 1.3^1	2.5 ± 1.2	G0.9 \pm 4.6	7.39 ± 0.4			
10669568	-0.1 ± 2.0^1	-10.9 ± 3.4	K4.2 \pm 0.6	5.68 ± 0.06	30.1	60.48 ± 0.17	0.01
10669590	-10.6 ± 3.5^1	-11.4 ± 2.3	K4.5 \pm 0.35	5.91 ± 0.04			
11017620*	-9.4 ± 2.1^1	-19.0 ± 0.5	F7.1 \pm 1.95	6.53 ± 0.24	8.9	-18.93 ± 0.6	0.95
11017626	-4.8 ± 1.2^1	-17.5 ± 0.9	F8.8 \pm 4.2	7.21 ± 0.38			
11551404	17.6 ± 1.5^1	12.6 ± 1.1	K2.6 \pm 0.6	4.66 ± 0.08	55.7	30.61 ± 2.2	0.00
11551430	11.2 ± 1.7^1	20.6 ± 1.4	K0.5 \pm 1.05	4.88 ± 0.13			
12253474	-11.4 ± 2.3^1	-8.1 ± 0.5	G0.7 \pm 4.65	7.7 ± 0.4	18.0	-1.02 ± 0.17	0.45
12253481	-5.4 ± 2.2^1	-7.0 ± 1.0	A6.6 \pm 5.05	8.23 ± 0.61			

This process is outlined in Appendix C. We performed our SED fits using the same process as described by Kraus & Hillenbrand (2007) fitting a reduced χ^2 for all good data points, removing the most discrepant data point around the best fit (the fit with the lowest reduced χ^2) if it is more discrepant than 3σ , refitting, redoing our clipping of the most discrepant point with the same conditions as before and finally refitting. We do not clip if it would leave fewer than four good data points and we not do an initial fit if we have fewer than this number. Note that for the Pan-STARRS1 photometry we exclude filters where the object is marked as extended and has fewer than four detections. We also exclude data from filters above the following saturation magnitude limits: $g_{P1} = 14.5$, $r_{P1} = 14.5$, $i_{P1} = 14.5$,

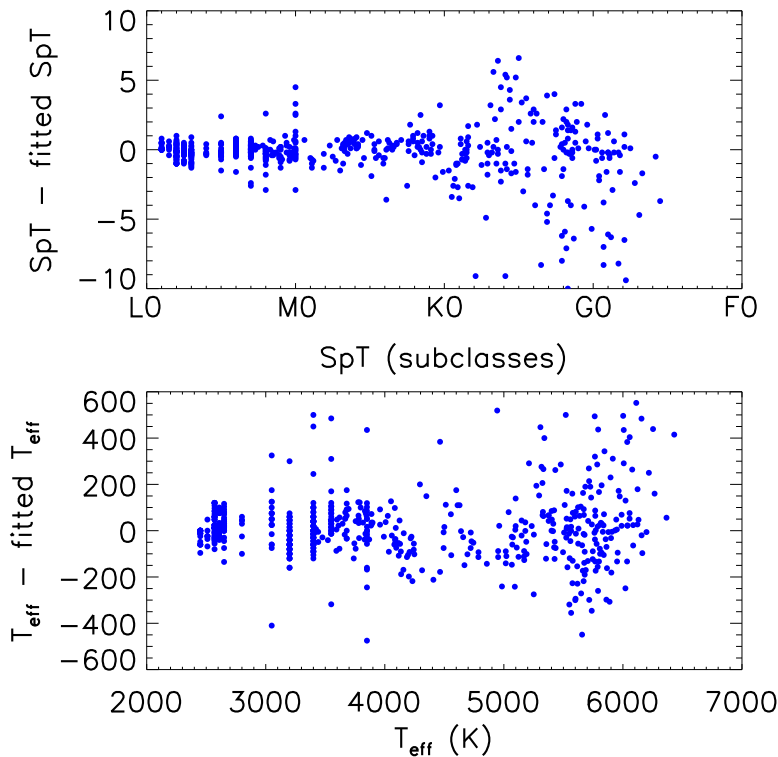


Figure 8. A comparison between the measured spectral types and effective temperatures for our calibration stars and those derived from fitting. Note the objects shown here have at least four photometric measurement and exclude out M9.5 and L0 calibration objects as we do not have an L0 template to fit these against. In the upper plot a positive y-axis value indicates that our spectral type estimate from SED fitting is earlier than that derived from spectroscopy. Many of the M dwarfs have effective temperatures derived from their spectral types, hence the discrete temperatures in this region of the plot

$z_{P1} = 14.5$, $y_{P1} = 14.5$ and $J_{MKO} = 10.5$. This last saturation limit is the same as that used in Deacon et al. (2009). It should be stated that our SED templates (like those used in Kraus et al. 2014b) assume the target is a dwarf. Hence this method may miss or incorrectly match binaries containing one or more evolved stars. To test the accuracy of our SED fitting method, we ran our calibration sample SEDs (as defined in Appendix B) through our fitting procedure. The results of these fits are shown in Figure 8. This plot only includes astronomical objects with actual Pan-STARRS 1 measurement used for calibration not the synthetic Pan-STARRS 1 magnitudes used to model hot stars.

The method produces reliable results for a range of spectral types, but with larger scatter for F and G type stars than K and M. Using a robust estimate of the scatter based on median absolute deviation we estimate that the standard deviation for our method is 2.5 subclasses for F and G stars and 0.6 subclasses for K and M stars. This equates to a scatter of 190 K and 75 K respectively. Kraus et al. (2014b) note that their SED fitting is accurate to less than one subclass. There is also the fact that the underlying calibration objects have measurement errors on their effective temperatures and/or those temperatures are derived from spectral

types which are typically quoted in discrete steps of half a subclass. For example our F and G calibration stars which are mostly drawn from Huber et al. (2014) have typical errors on their measured effective temperatures of about 110 K equating to 1.5 subclasses. This means that our quoted scatter in our measurements is likely to be around 180 K or 2.0 subclasses for F and G stars. We choose not to alter our scatter on K and M stars as we do not have a formal error on spectral type (and hence effective temperature) for many of these calibration objects. Our SED fitting procedure produces error estimates based on the calculated reduced χ^2 distribution. However these are often extremely small confidence ranges compared to the scatter we have derived here. Hence when an object has a confidence range in spectral type that falls below our calculated empirical scatter, we substitute our empirical scatter measurements to estimate one sigma confidence regions for our calculated distance moduli and bolometric magnitude. This results in a typical uncertainty in bolometric magnitude of 0.4 mag.

We included astrometry and SED fits for stars brighter than our Pan-STARRS 1 saturation limit of $r_{P1} < 14.5$ by including astrometry from either the UCAC4 catalogue (Zacharias et al. 2013) or the PPMXL catalogue (Roeser et al. 2010). Where stars had an entry in both catalogues we selected the data source with the lowest quoted error on proper motion. To avoid losing objects around our saturation boundary we did not set any limits on magnitude, we simply excluded UCAC4 or PPMXL measurements of any object with an unsaturated Pan-STARRS 1 proper motion measurement in our catalogue. We then added SED fits from Kraus et al. (2014b) giving us proper motion, spectral type and distance information for most bright stars in the *Kepler* field.

We used our combined catalogue as a basis for our wide binary selection. From it we selected possible companions to *Kepler* targets separated by less than five arcminutes from their primary and with distance moduli which differed by less than two magnitudes. We then ran our binary probability algorithm using our quoted errors on distance modulus and proper motion for all our possible pairings. A histogram of our membership probabilities is shown in Figure 9. We have a large number of low probability pairings and an excess of high probability pairings with $p_{bin} > 0.8$.

4.3.1 Exclusion of giants

One underlying assumption in our SED fits is that all the stars for which we derive distance moduli are dwarfs. This is clearly untrue as many stars in the *Kepler* field will be giants. Mann et al. (2012) showed that bright, red targets in the *Kepler* field are typically giants. To remove giants from our binary sample we set two cuts, one using g , r and D_{51} photometry from the *Kepler* Input Catalogue (Brown et al. 2011) and one using 2MASS photometry. Both these use the separation Mann et al. (2012) showed between the late-type giant and dwarf populations. We define giants as either having $J - K_s > 1$ or having $g - D_{51} > 0.25$ and $g - r < 1.3$. The first cut used here would exclude very late-type companion to a *Kepler* target so we do not impose this 2MASS related cut on our secondary stars. As an additional check we looked at the *Hipparcos* parallaxes (van Leeuwen 2007) of stars in our binary sample. Of the 9 stars with measured *Hipparcos* parallaxes, none had distances which were significantly discrepant from our photometric distance estimates.

We have endeavoured to minimise giant contamination but as the distance moduli for these binaries use the assumption that they are dwarfs, anyone using individual binaries from this sample should be aware that giant contamination may not have been completely eradicated and some giants may remain in the sample. We urge particular caution for the systems with primaries KIC 6367993 and KIC 10592818 as these have distance moduli less than one.

5 THE WIDE BINARY POPULATION IN THE KEPLER FIELD

After excluding giant stars and removing binaries closer than $6''^2$ we were left with Kepler targets in 401 binary systems which have binary probabilities above 80%. We list these in Table 4 and show their projected physical and angular separations in Figure 10.

5.1 The ages of wide binary components

Table 4 also contains rotation period estimates from McQuillan et al. (2014) for stars where this was measured. For pairs with two *Kepler* Targets with measured rotation periods and separations greater than $6''$ we estimated the ages of both components using the age-rotation relations of Mamajek & Hillenbrand (2008). Rather than using measured $B - V$

² This cut was to remove binaries where the presence of a close companion star would cause errors in the determination of exoplanet parameters and rotation periods.

values (which do not exist for most *Kepler* targets) for each target we converted our measured spectral types from SED fitting to colours using the field dwarf relation derived by Pecaú & Mamajek (2013). Figure 11 shows the ages derived for our pairs. Most have reasonable agreement between their ages. Four of our 7 binaries agree within one σ (roughly as expected) with one further binary having ages within two σ . We would expect from our $>80\%$ probability binary selection criteria that one or two of our seven proposed binaries would be not true physical pairs. Additionally, Mamajek & Hillenbrand (2008) find that four of the solar-type binaries in their test sample of 17 systems have ages which are discrepant by more than 0.3 dex. Hence our gyrochronology results do not indicate a larger than expected contamination in our binary pairs sample.

Table 4. Our candidate binary systems, the full table will be available electronically. Rotations periods from McQuillan et al. (2014) Key: ^a proper motion from Zacharias et al. (2013), ^bSED fit calculated using the methods of Kraus et al. (2014a), ^c proper motion/SED fit from this work, ^d proper motion from Roeser et al. (2010) .

KIC	KOI	R.A. (Ep=2013.0, Eq=J2000)	Dec.	$\mu_\alpha \cos \delta$ (mas/yr)	μ_δ (mas/yr)	Distance modulus (mag)	SpT (subtypes)	P_{rot} (days)	r (arcsec)	(AU)	P_{bin} Designation	WDS
1161145		19:23:59.11	+36:52:31.2	-5.8 ± 1.4^a	-3.5 ± 1.2	7.26 ± 0.07	G6.7 ^b		10.7	3021	0.81	WDS 19240+3653B
1161137		19:23:58.32	+36:52:26.2	-0.1 ± 1.3^a	-0.4 ± 1.6	7.22 ± 0.07	G7.1 ^b					WDS 19240+3653A
1295737		19:26:30.04	+36:54:47.7	5.2 ± 1.0^a	-5.8 ± 2.2	8.36 ± 0.21	F4.0 ^b		13.1	6166	0.84	
		19:26:30.99	+36:54:41.2	0.7 ± 1.8^c	-2.5 ± 1.6	8.35 ± 0.47	M2.3 ^c					
1571717		19:24:10.52	+37:06:34.4	-0.9 ± 1.9^a	2.1 ± 0.9	9.03 ± 0.36	A6.9 ^b		9.2	5868	0.82	
1571732		19:24:11.27	+37:06:32.3	-0.8 ± 3.7^d	2.4 ± 3.7	8.95 ± 0.4	F5.8 ^b					
1873370		19:30:12.45	+37:20:19.3	2.6 ± 2.1^c	-3.1 ± 2.1	10.65 ± 0.22	F9.1 ^c		7.0	9450	0.84	
1867358		19:30:12.24	+37:20:25.9	2.3 ± 2.0^c	-0.0 ± 1.8	10.81 ± 0.15	K7.6 ^c	32.879 ± 0.735				

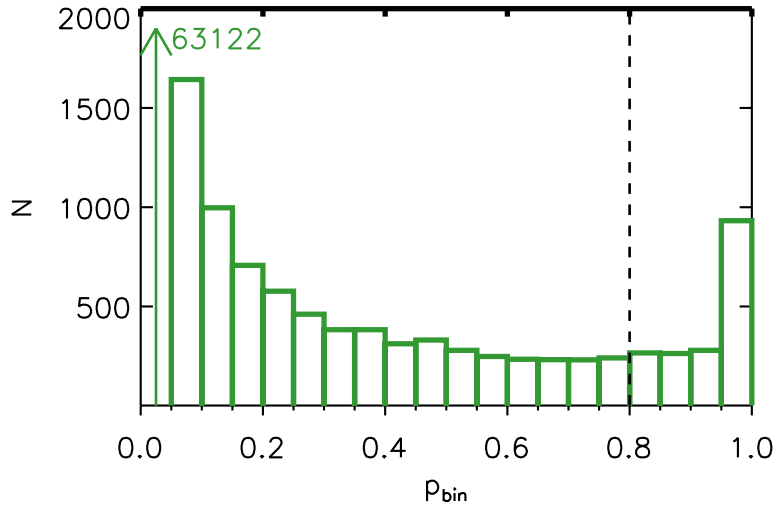


Figure 9. The distribution of binary probabilities for wide binaries with at least one *Kepler* target star. The arrow indicates the very large number of low probability pairings. The dashed line shows our 80% threshold for high probability binaries

5.2 The planet-host fraction of wide binary stars

To study how the presence of a wide binary companion affects the resulting planetary system we use our sample of wide binaries with companionship probabilities above 80% and which had passed our giant removal criteria. This sample of 401 multiples contained 529 *Kepler* targets. Next we crossmatched each *Kepler* targets in the sample with the most recent list of *Kepler* Objects of Interest (KOI) published by Rowe et al. (2015). These are divided into confirmed planets and planet candidates. If a binary system contained two *Kepler* target stars then we treated each *Kepler* target separately so a binary system where both *Kepler*

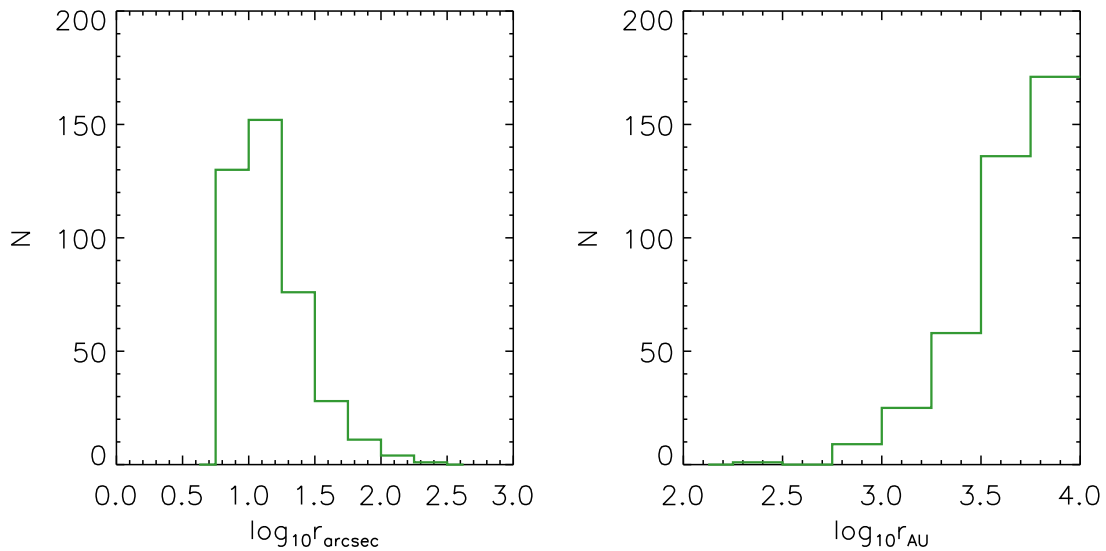


Figure 10. The properties of our binaries wider than 6'' with binary probabilities greater than 0.8. The separations are shown in arcseconds (left) and AU at the calculated distance of the *Kepler* target (right).

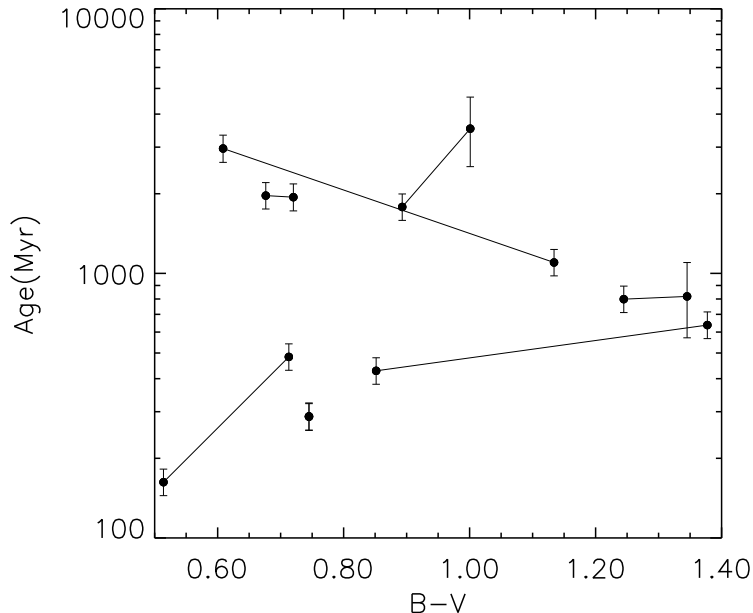


Figure 11. Ages calculated from gyrochronology for binaries with two periodic *Kepler* targets as components. The lines join the two components of each pair with horizontal lines representing good agreement between the ages of both components. The error bars on the ages include the 0.05 dex errors suggested by Mamajek & Hillenbrand (2008). What appears to be a single data point is the binary pair KIC 10913758 and KIC 10913762 which our SED fitting classifies as having the same spectral type and which have near-identical rotation periods.

targets were classified as a KOI we counted each of these as a separate planet-hosts. We restricted our sample to *Kepler* targets with $10 < K_p < 16$ and $0.5M_\odot < M < 1.5M_\odot$ as there are few *Kepler* target stars outside these ranges (where we assume a mass from our SED fitting). Our separation greater than $6''$ cut meant we largely sample binaries with projected separations above 3,000 AU (see Figure 10). After our restrictions on mass and K_p magnitude we were left with 419 objects from which we identified 10 KOI candidate planet hosts and 3 KOI confirmed planet hosts³ (with two additional candidate planet hosts with masses outside our $0.5M_\odot < M < 1.5M_\odot$ range). Hence $3.1 \pm 0.9\%$ of our likely binary primaries are candidate or confirmed planet hosts. No binary system where both components were KOIs was identified.

We also constructed a comparison sample of *Kepler* targets. These were 159152 objects which passed our giant exclusion criteria and had $10 < K_p < 16$ and $0.5M_\odot < M < 1.5M_\odot$. We found that our binary primaries were generally brighter and lower mass than the comparison sample of *Kepler* targets in general. This is likely due to our method selecting

³ We use the *NExTSci Disposition* for the candidate and confirmed planet host classifications

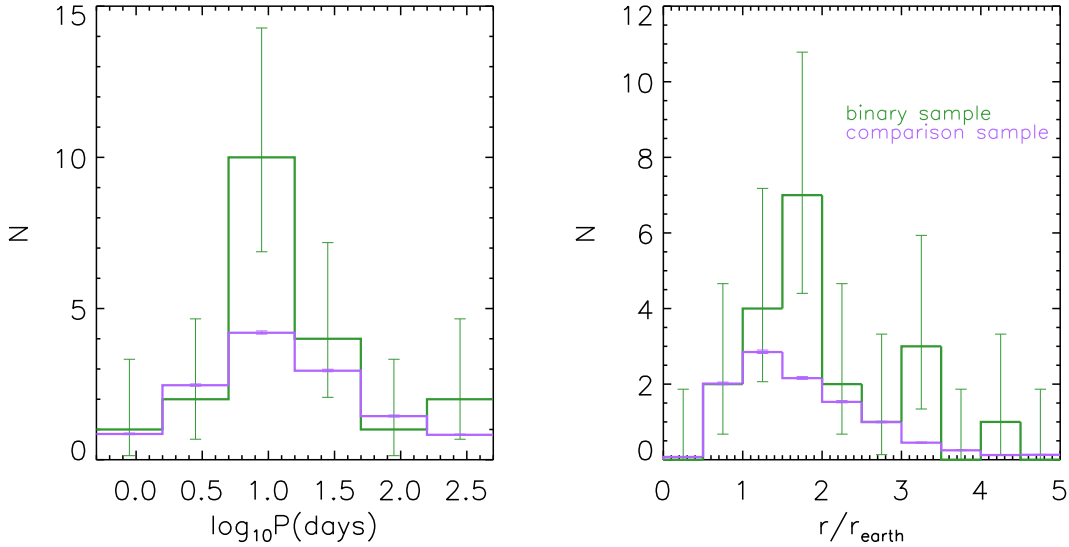


Figure 12. The properties of confirmed/candidate exoplanets around our binary components compared to a similar sample of *Kepler* targets. Our binaries exclude systems with separations below $6''$.

low-mass, nearby, high proper motion binaries preferentially as the number of coincident pairings with low-mass, nearby, high proper motion stars will be very low. To remedy this, we constructed 2D histograms for both our binary primaries and our comparison stars with bin sizes of 1 mag in K_p and $0.2M_\odot$ in mass. We then used these to construct a weighting factor for each comparison star based on its K_p magnitude and mass,

$$W(K_p, m) = \frac{N_{bin}(K_p, m)}{N_{comp}(K_p, m)} \quad (1)$$

where W is our weighting factor, N_{comp} is the number of comparison stars in that bin and N_{bin} is the number of binary primaries in that bin. The median bin in N_{bin} contained 12 objects while the median bin in N_{comp} contained 1804 objects. To estimate the number of planet hosts in our sample we crossmatched our comparison sample with the list of KOIs and then summed the weighting factors to calculate how many planet hosts would be expected in a sample of comparison stars of identical size, magnitude distribution and mass distribution. This led us to a comparison KOI rate of $2.4 \pm 0.5\%$. Hence we find no evidence for suppression or enhancement of the number of wide binary components which host exoplanets detectable by *Kepler*. We found that $1.0 \pm 0.5\%$ of our binaries had more than one *Kepler* confirmed or candidate planet while $0.5 \pm 0.1\%$ of our weighted sample did. This is again a statistically insignificant difference. We also examined the distributions of orbital period and planetary radius (see Figure 12) finding no statistically significant differences between the candidate and confirmed exoplanets around our binary components and around a similar sample of *Kepler* targets.

Our work finds no statistically significant difference between the fraction of wide binary components which are identified as being Kepler Objects of Interest and a sample of Kepler targets of similar brightness and spectral type. While these KOIs are not all confirmed planet hosts, the fact that there is no significant difference indicates that in the regime where both our binary and comparison sample have detect planets ($P < 300$ days) there is no evidence that a wide stellar companion with a separation wider than 3,000 AU will have a large effect on planet occurrence. Note that we do not claim that our comparison sample only contains single stars as many will have unresolved companions or companions which are too faint. This sample relies on a small number of binaries containing 419 Kepler targets. The ongoing *Kepler K2* mission will drastically increase the number of stars in wide binaries probed for transiting exoplanets. This will increase the accuracy of the measurement of any possible effect from wide binarity. Additionally *K2* will focus more on M dwarfs, objects largely excluded from our analysis here. Finally it should be noted that our study does not disprove the theoretical hypothesis of Kaib et al. (2013) that wide binaries can affect the orbits of planets. This effect could still be extremely significant for planets in wider orbits similar to the gas giants in our own Solar System.

6 CONCLUSIONS

We have used Pan-STARRS1 data and archival datasets to calculate accurate proper motions for stars in the *Kepler* field. This catalogue will be available to the community online and will improve the characterisation of Kepler target stars (i.e. giant-dwarf separation, kinematic population membership etc.). By combining the proper motions with PS1-based SEDs to estimate the distance, we selected a sample wide binaries. To formalise this, we developed a statistical method to calculate true association probabilities for wide binaries including empirical estimates of the frequency of pairings of unrelated field stars. Our method was tested for a sample of bright stars showing that our likely binaries had low differences in radial velocity while our unlikely pairings had discrepant radial velocities. After excluding giant stars, we find 401 multiples containing at least one *Kepler* target which have binary companion probabilities of 0.8 or more and separations between six arcseconds and five arcminutes. We find no difference in the rate at which our wide binaries (typically wider than 3,000 AU) are identified as confirmed or candidate exoplanet hosts than a comparison sample of *Kepler* targets of similar brightness and spectral type. An investigation of those

binaries where both components have archival rotation periods shows that our binaries are roughly coeval indicating that our method is selecting true binaries.

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APPENDIX A: DETERMINING COMPANIONSHIP PROBABILITIES

In order to define a formal companionship probability, we define the following likelihood,

$$\phi = \phi_c(\Delta\mu, \mu, \Delta D, D, s, m) + \phi_f(\mu, D, s, m) \quad (\text{A1})$$

where ϕ_c and ϕ_f are the companion and coincident field distributions and are functions for the separation (s), proper motion difference ($\Delta\mu$), proper motion (μ) and distance modulus difference (ΔD) between any particular pair and the mass of the secondary component (m). The probability of a pair being real rather than a coincident pairing of unrelated field stars is

$$p = \frac{\phi_c}{\phi_c + \phi_f} \quad (\text{A2})$$

First let us examine the companion distribution,

$$\phi_c = f_{bin} f_{det} \times \left[\frac{e^{-\Delta\mu^2/2\sigma_\mu^2}}{2\pi\sigma_\mu^2} \right] \times \left[\frac{e^{-\Delta D^2/2\sigma_D^2}}{\sqrt{2\pi}\sigma_D} \right] \times [c_s (sd)^{-x}] \times [c_m f(m)] \quad (\text{A3})$$

Where $\Delta\mu$ is the proper motion difference, σ_μ is the quadrature sum of the proper motion errors, ΔD is the difference in distance moduli, σ_D is the quadrature sum of the distance modulus errors, s is the separation on the sky, m is the mass of the secondary and d is the calculated distance to the primary. The last two sets of brackets on the right each contain normalisation terms. We assume a power-law in physical separation but this method could be adapted to any underlying separation distribution or binary fraction. The separation distribution is normalised by c_s ,

$$c_s = 1 / \int_{s_1}^{s_2} (sd)^{-x} ds \quad (\text{A4})$$

Where s_1 and s_2 are the inner and outer search radii of our survey. There are two free parameters to choose to complete this system, the first is the exponent x which defines the separation distribution. The other is the binary fraction f_{bin} defined as the fraction of stars which have a companion in the region in which our separation distribution is defined. Next we must consider our observational selection effect caused by our choice of an inner and outer search separation. This is written in Equation A3 as f_{det} and can be defined as,

$$f_{det} = d \int_{s_1}^{s_2} (sd)^{-x} ds / \int_{r_{min}}^{r_{max}} r^{-x} dr \quad (\text{A5})$$

Where r is the physical separation in AU (defined as sd) and our separation distribution is defined in the region r_{min} to r_{max} . Note the numerator of Equation A5 is equal to the denominator of Equation A4. Hence we can condense these two factors so that,

$$c'_s = c_s \times f_{det} = d / \int_{r_{min}}^{r_{max}} r^{-x} dr \quad (\text{A6})$$

The final term in Equation A3 is the mass distribution of likely companions ($f(m)$) with the normalisation c_m . For this study we assume a flat mass ratio distribution. Hence as the mass of the secondary cannot be greater than the mass of the primary, the normalisation c_m becomes $1/m_p$ where m_p is the mass of the primary (assumed to be the earlier-type star in the pair).

$$\phi_c = f_{bin} \times \left[\frac{e^{-\Delta\mu^2/2\sigma_\mu^2}}{2\pi\sigma_\mu^2} \right] \times \left[\frac{e^{-\Delta D^2/2\sigma_D^2}}{\sqrt{2\pi}\sigma_D} \right] \times [c'_s (sd)^{-x}] \times \left[\frac{1}{m_p} \right] \quad (\text{A7})$$

Next let us turn to the field distribution. The Galactic field star population is complex with the proper motion distribution particularly hard to parameterise over a wide range of proper motions. Hence we adapt the positional offset test of Lépine & Bongiorno (2007) to measure the likely contamination rate. Defining,

$$\phi_f = c_f 2\pi s \tag{A8}$$

where c_f is the density of stars of similar absolute proper motion and has units of stars per (mas/yr.)² per magnitude per sq. arcsecond per solar mass. The factor c_f is defined empirically by offsetting the positions of primary stars in our sample and repairing them to produce entirely confident pairs (Lépine & Bongiorno 2007). We do this and (for each of our actual pairs) measure the density of coincident pairs at the particular magnitude difference, proper motion difference and mass ratio of the actual pair.

APPENDIX B: EMPIRICAL SEDS IN THE PAN-STARRS 1 PHOTOMETRIC SYSTEM

In order to calculate accurate photometric distances for stars in the *Kepler* field we require a series of empirical stellar Spectral Energy Distributions (SEDs). There are multiple sources of empirical stellar SEDs with the most commonly used being Kraus & Hillenbrand (2007) for the SDSS/2MASS systems and Pecaut & Mamajek (2013) for the Cousins/2MASS/WISE systems. There are no empirical SEDs in the Pan-STARRS 1 photometric system defined by Tonry et al. (2012) over the required range of spectral types (A–M). Dupuy & Liu (2009) present synthetic colours for a range of spectral types based on initial characterisation of the throughput of the Pan-STARRS 1 photometric system. Aller et al. (2013) present a series of different SED templates for M dwarfs. These include templates for M and early L dwarfs in the star forming region Upper Sco and early M dwarfs based on members of the Praesepe and Coma Ber clusters. Finally Aller et al. (2013) presents individual M9 and L0 dwarfs to be used templates for field stars. Note that Aller et al. (2013) used photometry from a previous processing version to our work so there may be subtle differences between the photometric systems.

We began by taking the large collection of known dwarfs in the *Kepler* field with spectrally determined effective temperature from Huber et al. (2014) and references therein. We converted their effective temperatures to spectral type using the dwarf effective temperature scale of Pecaut & Mamajek (2013). For M dwarfs from the work of Martin et al. (2013)

we note that the effective temperatures in this work are derived from SED fitting so instead we convert the objects' determined spectral types to effective temperatures using the Pecaut & Mamajek (2013) scale. We supplement the objects from the Huber et al. (2014) collection with late M dwarfs drawn from the compilation of Faherty et al. (2009) as these objects typically have high proper motions we corrected their positions to an epoch 2010.0 for comparing to Pan-STARRS 1 data. We also added M dwarfs from the sample of West et al. (2008) taking only bright ($z_{sdss} < 16$) and low extinction ($A_g < 0.05$) objects. We drew Pan-STARRS 1 photometry from the nightly science database and 2MASS data from the 2MASS catalogue using pairing radii of 3 arcseconds. Objects were only incorporated into the fit in a particular PS1 band if they had a magnitude error less than 0.05, more than 2 detections, were not saturated in that band, were not extended in that band and had been subject to the "ubercal" calibration process (Schlafly et al. 2012).

We note that there are few objects in the Huber et al. (2014) catalogue hotter than 6000 K which are not saturated in at least one Pan-STARRS 1 band. Hence we supplemented our objects with synthetic colours calculated from Pickles (1998) and Dwarf Archives optical spectra and near-IR spectra from Cushing et al. (2005) and Rayner et al. (2009). The Pickles spectra overlapped with those from Rayner over the y_{P1} band so we calculated mean optical colours relative to y_{P1} in the AB system for each spectral type. We then calculated the colours relative to J_{2MASS} by calculating $y_{P1} - J_{2MASS}$ from the Rayner spectra, adding a term to take into account that y_{P1} is in the AB system and J_{2MASS} is in the Vega system.

For the Pan-STARRS 1 bands we developed the following mechanism to calculate the magnitudes for each spectral type in each band. We began by fitting our $g_{P1} - i_{P1}$ colours with a spline to produce a smooth distribution relative to spectral type. For spectral type we used ten subclasses per spectral classification. This is a departure from the traditional picture where there are no K6, K8 and K9 dwarfs. However we use the effective temperature scale of Pecaut & Mamajek (2013) who include these missing spectral types in their temperature scale. Our $g_{P1} - i_{P1}$ to spectral type fit is shown in Figure B1. We then produced splines for relations between $g_{P1} - i_{P1}$ and other colours; $r_{P1} - i_{P1}$, $i_{P1} - z_{P1}$, $z_{P1} - y_{P1}$ and $y_{P1} - J_{2MASS}$. These fits are shown in Figure B2. We then used these colours at each $g_{P1} - i_{P1}$ (and hence spectral type) to produce colours relative to the 2MASS J -band magnitude.

These were then combined with the bolometric magnitudes from Kraus & Hillenbrand (2007) and the bolometric correction and $V - J$ colours of Pecaut & Mamajek (2013). for objects later than M7 we used the spectral type - J -band absolute magnitude relations of

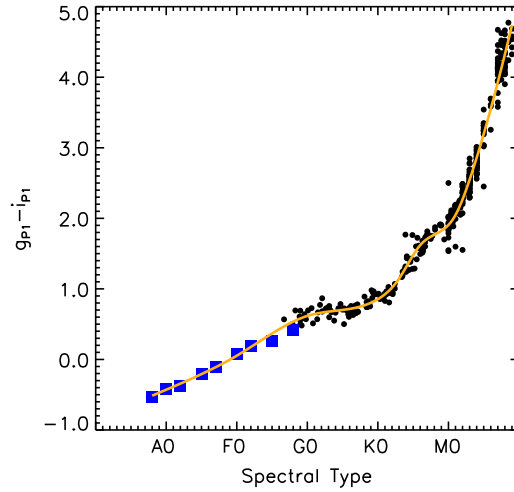


Figure B1. The spectral type to $g_{p1} - i_{p1}$ transformation used to calibrate our SED templates. The black points are observed stars of known temperature/spectral type and the blue squares are synthetic photometry derived from the spectra of Pickles (1998) and Rayner et al. (2009). The yellow line is a cubic spline fit to the data.

Dupuy & Liu (2012). For near-infrared colours, we used $J_{2MASS} - H_{2MASS}$ and $J_{2MASS} - K_{s2MASS}$ colours derived from the bolometric corrections of Pecaut & Mamajek (2013). We also synthesised $J_{2MASS} - J_{MKO}$, $H_{2MASS} - H_{MKO}$ and $K_{s2MASS} - K_{MKO}$ from the Rayner spectra to derive the MKO magnitudes (Tokunaga et al. 2002) of all our spectral types earlier than M7 and used the quoted Dupuy & Liu (2012) spectral type – absolute magnitude relations for spectral types M7 and later. Our input sample included a handful of M9.5 and L0 objects to anchor our fits at the cool end. We do not quote an L0 template. Our final SED templates for each subtype between B9 and M9 are shown in Table B1.

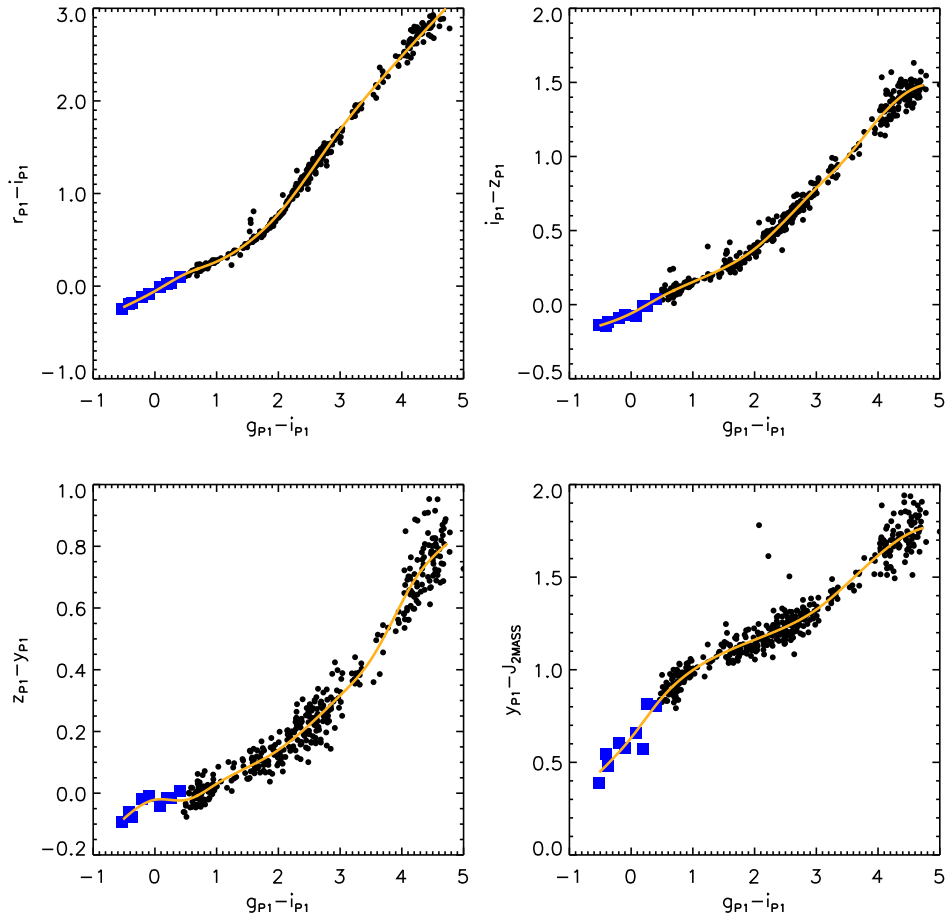


Figure B2. The transformation between $g_{P1} - i_{P1}$ and various other colours used to calibrate our SED templates. The black points are observed stars of known temperature/spectral type and the blue squares are synthetic photometry derived from the spectra of Pickles (1998) and Rayner et al. (2009). The yellow line is a cubic spline fit to the data.

Table B1. The SED templates in the Pan-STARRS1 and MKO photometric systems used in this work. The bolometric magnitudes and masses come from the compilation of Kraus & Hillenbrand (2007) and the effective temperatures are from the dwarf effective temperature scale is the dwarf scale from Pecaut & Mamajek (2013). Our templates also used the 2MASS colours from Pecaut & Mamajek (2013) before M6 and from Dupuy & Liu (2012). The reader should note that Pecaut & Mamajek (2013) and Dupuy & Liu (2012) also include WISE colours.

SpT	g_{P1} (mag.)	r_{P1} (mag.)	i_{P1} (mag.)	z_{P1} (mag.)	y_{P1} (mag.)	J_{MKO} (mag.)	H_{MKO} (mag.)	K_{MKO} (mag.)	M_{bol} (mag.)	T_{eff} K	mass M_{\odot}
B8.0V	-0.31	-0.03	0.2	0.34	0.43	0.02	0.06	0.16	-1.0	12500	3.8
B9.0V	-0.05	0.21	0.42	0.56	0.64	0.21	0.22	0.29	-0.35	10700	3.35
A0.0V	0.34	0.57	0.77	0.9	0.97	0.51	0.51	0.54	0.3	9700	2.9
A1.0V	0.67	0.87	1.06	1.18	1.24	0.78	0.76	0.79	0.7	9200	2.65
A2.0V	1.02	1.19	1.36	1.48	1.54	1.07	1.03	1.06	1.1	8840	2.4
A3.0V	1.24	1.38	1.54	1.65	1.7	1.21	1.17	1.2	1.32	8550	2.267
A4.0V	1.41	1.52	1.67	1.77	1.81	1.3	1.24	1.28	1.53	8270	2.133
A5.0V	1.65	1.73	1.86	1.96	2.0	1.46	1.41	1.44	1.75	8080	2.0
A6.0V	1.87	1.92	2.03	2.12	2.15	1.6	1.54	1.58	1.92	8000	1.9
A7.0V	2.02	2.03	2.13	2.21	2.23	1.67	1.59	1.63	2.08	7800	1.8
A8.0V	2.2	2.18	2.26	2.33	2.36	1.77	1.67	1.71	2.26	7500	1.733
A9.0V	2.46	2.41	2.46	2.53	2.55	1.93	1.83	1.86	2.43	7440	1.667
F0.0V	2.67	2.58	2.62	2.67	2.69	2.04	1.92	1.95	2.61	7200	1.6
F1.0V	2.84	2.72	2.73	2.76	2.78	2.12	1.97	1.98	2.75	7030	1.55
F2.0V	3.02	2.86	2.84	2.86	2.88	2.2	2.02	2.02	2.89	6810	1.5
F3.0V	3.35	3.15	3.11	3.11	3.14	2.42	2.23	2.23	3.13	6720	1.417
F4.0V	3.66	3.42	3.36	3.35	3.37	2.63	2.42	2.43	3.37	6640	1.333
F5.0V	3.96	3.68	3.59	3.57	3.59	2.83	2.6	2.61	3.61	6510	1.25
F6.0V	4.2	3.89	3.78	3.74	3.76	2.96	2.71	2.72	3.82	6340	1.223
F7.0V	4.48	4.12	3.99	3.94	3.96	3.13	2.88	2.88	4.03	6240	1.197
F8.0V	4.75	4.36	4.21	4.15	4.17	3.31	3.05	3.06	4.24	6150	1.17
F9.0V	4.91	4.49	4.33	4.26	4.27	3.4	3.13	3.13	4.35	6040	1.14
G0.0V	5.04	4.59	4.42	4.34	4.35	3.47	3.19	3.18	4.47	5920	1.11
G1.0V	5.13	4.66	4.49	4.4	4.41	3.52	3.22	3.21	4.53	5880	1.085
G2.0V	5.16	4.68	4.5	4.41	4.42	3.52	3.19	3.18	4.6	5770	1.06
G3.0V	5.27	4.78	4.59	4.5	4.51	3.61	3.27	3.26	4.7	5720	1.053
G4.0V	5.38	4.87	4.69	4.59	4.6	3.69	3.35	3.34	4.79	5680	1.047
G5.0V	5.48	4.97	4.78	4.68	4.69	3.78	3.44	3.42	4.89	5660	1.04
G6.0V	5.62	5.1	4.91	4.81	4.81	3.9	3.54	3.52	5.03	5590	1.02
G7.0V	5.79	5.25	5.05	4.95	4.95	4.03	3.67	3.65	5.16	5530	1.0
G8.0V	5.94	5.38	5.17	5.06	5.06	4.13	3.76	3.74	5.3	5490	0.98
G9.0V	6.17	5.58	5.37	5.25	5.24	4.3	3.91	3.89	5.49	5340	0.94
K0.0V	6.39	5.76	5.54	5.41	5.4	4.43	4.03	4.0	5.69	5280	0.9
K1.0V	6.67	5.99	5.75	5.61	5.59	4.61	4.18	4.15	5.89	5170	0.86
K2.0V	6.96	6.21	5.94	5.78	5.75	4.75	4.29	4.26	6.08	5040	0.82
K3.0V	7.32	6.48	6.17	5.99	5.94	4.92	4.41	4.37	6.32	4840	0.785
K4.0V	7.68	6.75	6.37	6.16	6.1	5.04	4.49	4.43	6.55	4620	0.75
K5.0V	8.09	7.07	6.62	6.37	6.29	5.2	4.63	4.56	6.68	4450	0.7
K6.0V	8.31	7.22	6.69	6.42	6.32	5.22	4.59	4.49	6.78	4200	0.665
K7.0V	8.51	7.38	6.8	6.5	6.4	5.29	4.64	4.52	6.89	4050	0.63
K8.0V	8.81	7.66	7.04	6.74	6.62	5.5	4.85	4.72	7.13	3970	0.617
K9.0V	9.11	7.94	7.29	6.96	6.84	5.71	5.07	4.92	7.36	3880	0.603
M0.0V	9.46	8.27	7.56	7.21	7.08	5.93	5.3	5.15	7.6	3850	0.59
M1.0V	9.99	8.76	7.94	7.55	7.4	6.24	5.64	5.45	7.97	3680	0.54
M2.0V	10.8	9.54	8.54	8.06	7.88	6.68	6.07	5.87	8.44	3550	0.42
M3.0V	11.9	10.61	9.35	8.76	8.53	7.28	6.7	6.48	9.09	3400	0.29
M4.0V	13.27	11.97	10.41	9.68	9.39	8.06	7.53	7.26	9.92	3200	0.2
M5.0V	15.03	13.69	11.84	10.97	10.61	9.19	8.65	8.35	11.01	3050	0.15
M6.0V	16.58	15.18	13.03	12.01	11.56	10.04	9.46	9.12	12.06	2800	0.12
M7.0V	17.88	16.38	13.95	12.73	12.14	10.56	10.01	9.57	12.7	2650	0.11
M8.0V	19.05	17.45	14.73	13.34	12.62	10.94	10.36	9.87	13.13	2570	0.102
M9.0V	20.05	18.33	15.31	13.83	13.02	11.21	10.59	10.09	13.43	2450	0.088

APPENDIX C: PAN-STARRS1 TO *KEPLER* PHOTOMETRIC TRANSFORMATIONS

Pan-STARRS1 provides CCD photometry for both the original *Kepler* field and the ten K2 mission fields. *Kepler* magnitudes are derived from a number of sources including Tycho Hog et al. (2000), photographic plates and g , r , i photometry in the *SDSS* system from the *Kepler* Spectroscopic Classification Program Brown et al. (2011). There are also additional datasets such as the *Kepler* INT Survey (KIS; Greiss et al. 2012). In order to enable the

Table C1. Color transformations from Pan-STARRS 1 filters to *Kepler* magnitudes. These take the form $y = \sum_{i=0}^4 c_i x^i$ with the applicable colour range for each fit listed as x_{range}

y_{color}	x_{color}	c_0	c_1	c_2	c_3	c_4	x_{range}
K_p-g_{P1}	$g_{P1}-r_{P1}$	0.0208986	-0.871090	-0.381325	1.50797	-1.12176	-0.28-1.46
K_p-r_{P1}	$r_{P1}-i_{P1}$	0.0502748	0.393252	-0.650769	0.126130	-0.00725137	-0.24-2.70
K_p-i_{P1}	$i_{P1}-z_{P1}$	0.131700	2.48210	-1.94587	0.572845	-0.0636664	-0.15-3.13
K_p-z_{P1}	$i_{P1}-z_{P1}$	0.131700	3.48210	-1.94587	0.572845	-0.0636664	-0.15-3.13

calculation of *Kepler* magnitudes from Pan-STARRS 1 data we synthesised *Kepler* magnitudes for a series of stars using the dwarf spectral sequence of Pickles (1998) for types B9 to M6. We supplemented these with seven late M dwarfs from the Dwarf Archives spectral library ⁴ with sufficient spectral coverage and three L and T dwarfs from the compilation of Leggett et al. (2010). Note for the Leggett et al. (2010) spectra we included objects with wavelength coverage until 4000Å. This leaves the 3800–4000Å region with no data in the *Kepler* filter bandpass, however for objects this cool the emission in this region will be negligible. We synthesised colours using the *Kepler* bandpass ⁵ and the Pan-STARRS 1 filter set (Tonry et al. 2012). All colours were calculated in the AB system. Our colours are shown in Figure C1 along with fifth degree polynomial fits. These coefficients of these fits are shown in Table C1. Figure C1 also shows the synthesised *Kepler* magnitudes for the *Kepler* standards listed in Brown et al. (2011) plotted against their observed Pan-STARRS1 magnitudes. Except for two outliers which lie near the saturation point for Pan-STARRS 1, there is clearly fairly good agreement between our synthesised *Kepler*-Pan-STARRS 1 and the Brown synthesised *Kepler*-real Pan-STARRS 1 colors. The small offsets are likely due to the differences in the ways the *Kepler* magnitudes were synthesised between the two sources.

⁴ <http://spider.ipac.caltech.edu/staff/davy/ARCHIVE/index.shtml>

⁵ http://keplergo.arc.nasa.gov/kepler_response_hires1.txt

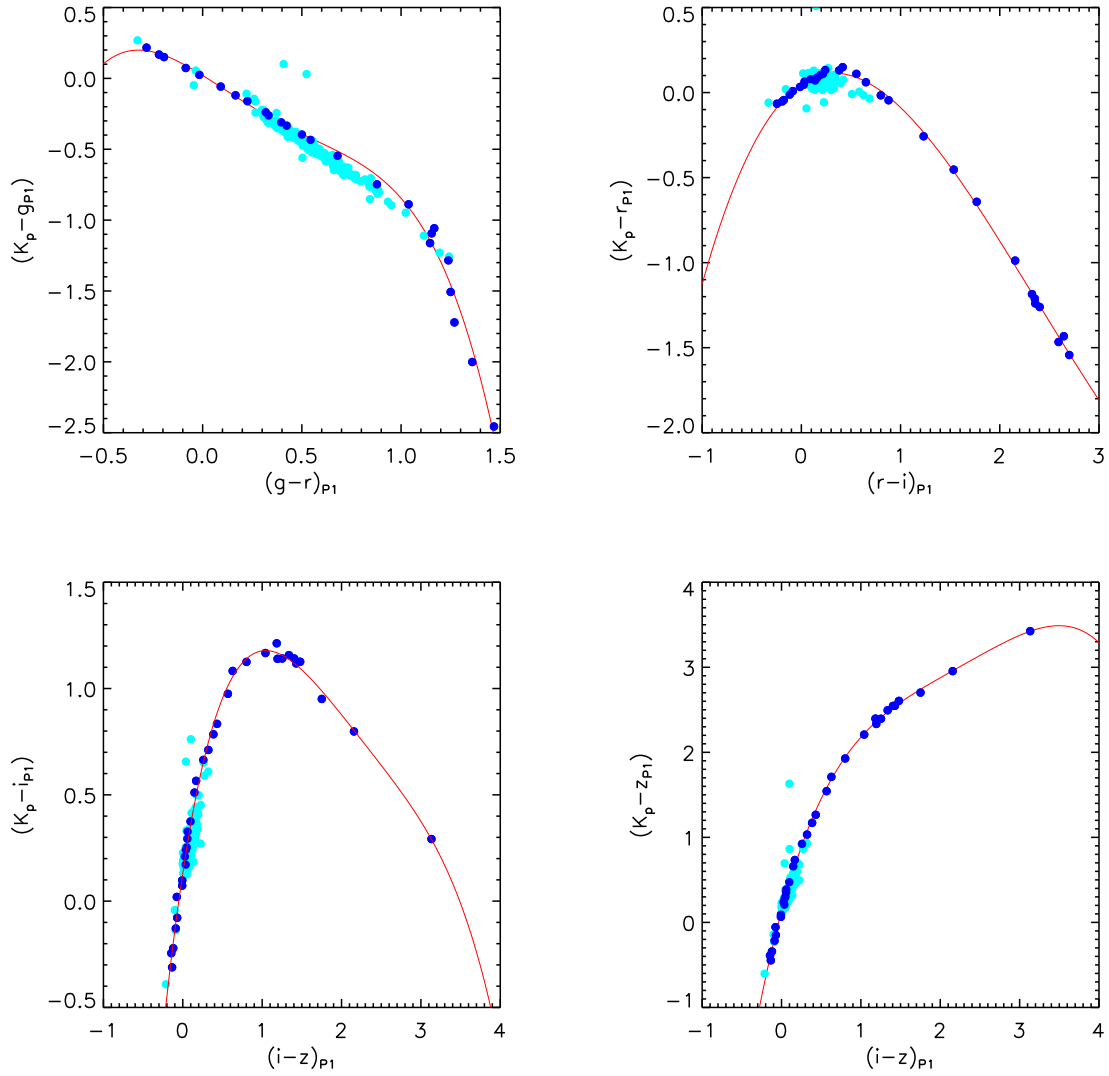


Figure C1. Synthesised colors for Pan-STARRS1 to *Kepler* magnitudes for empirical spectra from Pickles (1998), Dwarf Archives and Leggett et al. (2010) (blue points). The red line shows the fifth degree polynomial fits listed in Table C1. The aqua points are from the synthesised K_p magnitudes for the standards of Brown et al. (2011) and the observed Pan-STARRS1 magnitudes of these objects.