

COMMISSIONS G1 AND G4 OF THE IAU  
INFORMATION BULLETIN ON VARIABLE STARS

Volume 63 Number 6267 DOI: 10.22444/IBVS.6267

Konkoly Observatory  
Budapest

8 May 2019

HU ISSN 0374 – 0676

## NEW LIGHT ON R ARAE

BLANE, D.; BLACKFORD, M.G.; BUDDING, E.; REED, P.A.

Variable Star Section, Royal Astronomical Society of New Zealand, New Zealand

**Abstract**

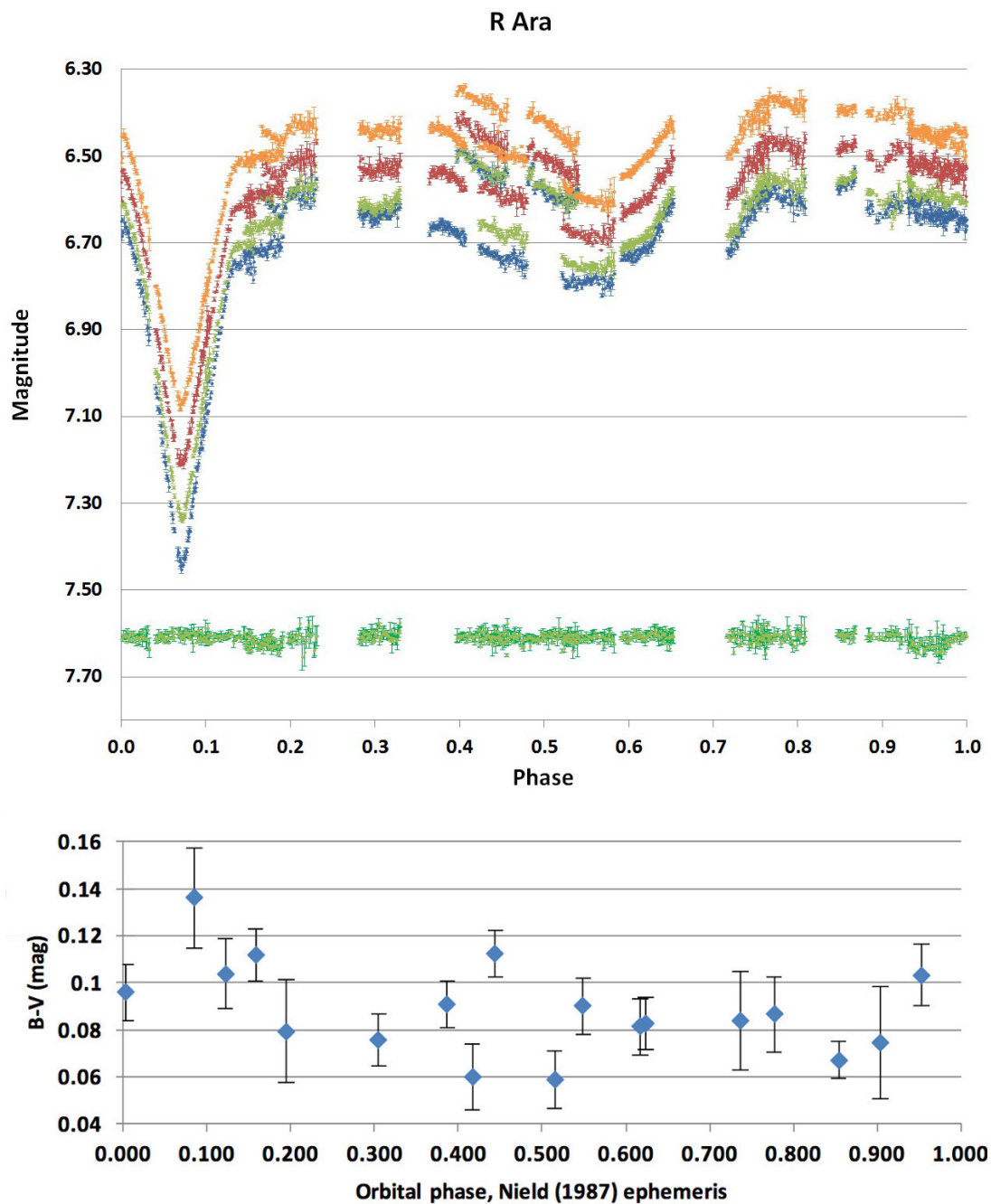
In mid-2018, efforts were renewed to check on the highly active mass-transfer process shown by the classical Algol system R Arae. We present new light curves and times of light minimum from this project. We also extend R Arae's period O–C diagram to include the new results. The new data are consistent with the strong mass transfer scenario of Reed (2011). Recent and ongoing studies of this interesting system are referred to.

## 1 Introduction

In the early 1980s investigators using satellite observational techniques were drawing attention to the relatively bright southern binary system R Arae, which superficially resembles a 'classical' Algol, but with mass transfer on a more enhanced scale than typical, and with additional photometric peculiarities that may be related to relatively dense and uneven accretion structures (Kondo et al., 1985; Nield et al., 1986; Reed, 2011). Kondo et al. (1985), perhaps with the symmetric, smoothly rounded light curve of Gaposchkin (1953) in mind, compared the system to the well-known massive and strongly interactive binary  $\beta$  Lyrae, despite the 3-times greater period of the latter.

The period steadily increased from about 4.42495 d in the year of its discovery by Roberts (1894) to the 4.425132 d given by Nield (1986; for HJD 2446585.161) and continues to do so according to Reed (2011), who obtained a mean rate of increase of  $5.15 \times 10^{-9} \text{ d d}^{-1}$  over 116 y. Reed deduced, from this period extension, that the Roche-Lobe filling component in the binary was shedding matter to its companion at a rate of about  $3.06 \times 10^{-7} M_{\odot} \text{ y}^{-1}$ .

Reliable parametrization of R Arae has been compromised by the significant ( $\sim 10\%$ ) level of short term variability that adds into the light curve (Banks, 1990). Such variations can occur even in one night's observations, as noted by Forbes et al. (1988), Budding (1989), and confirmed in Bakiş et al.'s (2016) study that included a short term photometric sequence from the HIPPARCOS satellite.

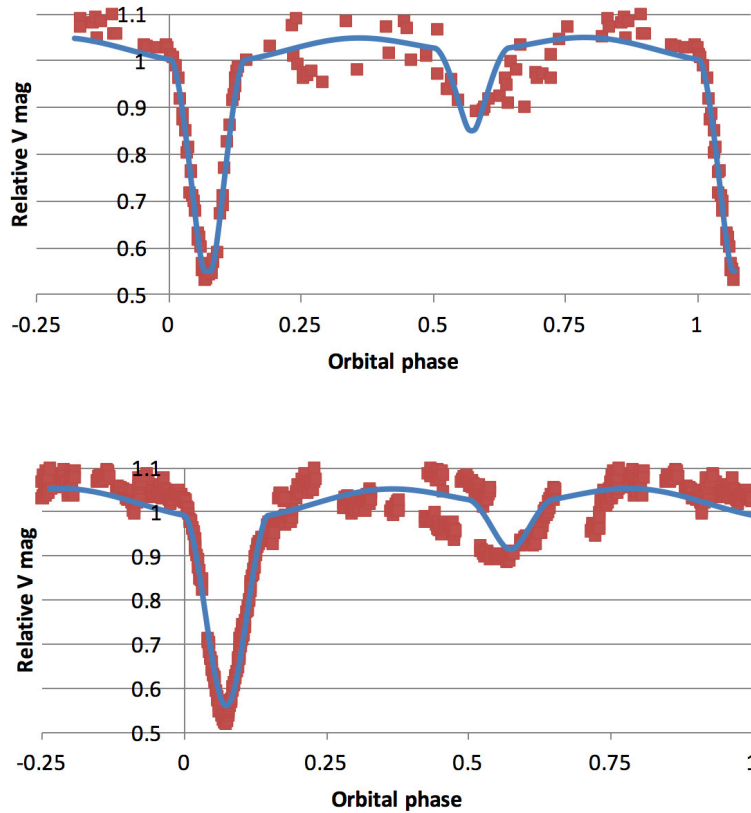


**Figure 1.** *BVRI* light curves and  $B - V$  colour curve of R Arae in mid-2018. The epoch and period from Nield (1987) were used for the phase calculation, see also Reed (2011).

## 2 New data

Light curves of R Arae were produced in Jun 2018 (DB) and Jul-Aug 2018 (MGB) as contributions to the Southern Binaries Programme of the VSS, stimulated by a new data request (PAR). For the June data, images were obtained with a 150 mm f5 refractor and a Canon 1300D DSLR camera mounted on a GEM goto mounting. Aperture photometry was facilitated using the IRIS software package. The camera's  $G$  magnitudes were linearly transformed to the Johnson  $V$  band using MS-Excel. Differential extinction was not applied. Each adopted measurement was the average obtained from 10 separate field images.

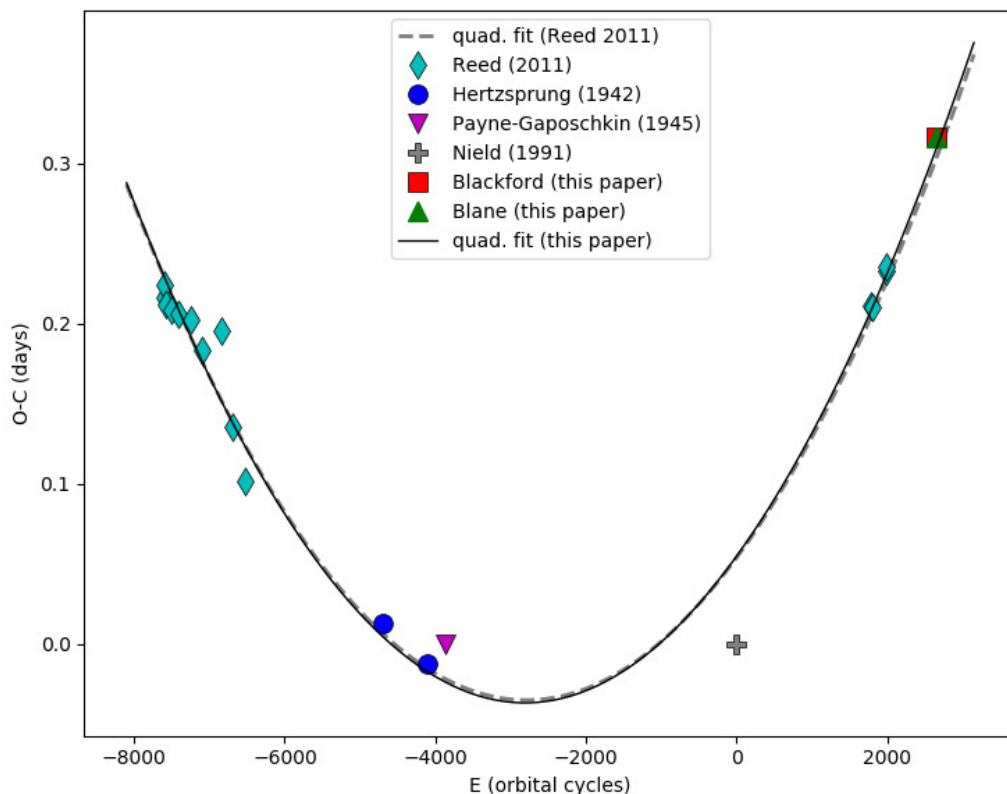
The Jul 7–Aug 19 data were gathered over 19 nights using an 80 mm f6 refractor and SBIG STT 3200ME CCD camera with Astrodon filters. These observations were extinction corrected and linearly transformed to standard Johnson-Cousins  $BVR_CI_C$  system. In Fig. 1 we present these latter data (upper panel: blue points  $B$ , green  $V$ , red  $R_C$ , orange  $I_C$ ). The lower green points show the check star, HD149519, shifted to an average value of 7.6 for easy display. Its actual measured average  $V$  value was 8.547 mag. The (binned)  $B - V$  colour curve is shown in the lower panel. The 19 separate runs consist of typically  $\sim 50$  points. The light-curves show the same kind of previously reported behaviour, with some features having a degree of persistence, for example the rise and decline into the primary eclipse, while other transient occurrences have a stochastic quality (Forbes, 1988).



**Figure 2.** The  $V$  light curves of 2018 Jun (DB, upper panel) and Jul-Aug (MB, lower panel), fitted with a classical Algol model. The phase shifts noted below are clearly visible in the light curves.

In the  $B - V$  colour curve (lower panel) we have binned each of the 18 sets of data into single average points, together with their s.d. dispersions. It is clear that the vagaries of colour are significantly greater than the standard deviation within the groups. The system is redder at mid-occultation, however – by about 0.06 mag. There is a suggestion of some blueing around the transit, but that is not uniform. The average  $B - V$  in the out-of-eclipse regions is 0.081 mag. These irregularities mean that parameters associated with fitting the light curves with standard models are not well-defined. Even so, a classical Algol model will approximately fit, as shown in Fig. 2, though other possibilities cannot be ruled out from photometric evidence alone. The significant secondary minimum and relatively low depth of the primary might be associated with a detached pair; though, when the light from the close companion (included in aperture photometry) is taken into account, the depth of the primary eclipse implies a secondary size comparable to its surrounding Roche lobe (Nield, 1987).

Bakiş et al. (2016), posited a model of cyclic phenomena in their account, though the question of why certain Algols show variable light curves (Piotrowski et al., 1974) and not all was more squarely addressed in the review of Reed (2012), who noted that certain period and mass-ratio combinations are propitious in allowing the development of relatively large and unsteady accretion structures. This model is in keeping with the relatively large rate of period increase demonstrated by Reed (2011) and confirmed in the latest (2018) times of minima illustrated in Fig. 3.



**Figure 3.** O–C variation as predicted by Reed (2011) and confirmed by our latest times of minima.

Times of primary minimum have been calculated as HJD 2458307.6501 (DB) and HJD 2458312.0758 (MGB), which give O–C values of 0.3157 d at cycle number 2649, and 0.3163

d at cycle 2650 from Nield's epoch.

It is of interest to note that relevant precise photometric data from NASA's Transiting Exoplanet Survey Satellite (TESS) should become available from Sector 12, that is expected to run between May 21 and June 19 this year. The foregoing  $BVR_CI_C$  information, not available from TESS, should then have a useful complementary role. Other complementary data on R Arae is available from KELT (Collins et al., 2018), together with spectroscopic data planned from MINERVA (Wittenmyer et al., 2018) or contained in the University of Canterbury Mt John Observatory HERCULES archive. This present report will thus hopefully contribute to ongoing observational work on this intriguing active close binary system.

*Acknowledgements* We appreciate the helpful suggestions of Dr L. Molnár, editor of the IBVS, regarding this contribution, and keenly appreciate the support that the IBVS has given to this kind of research over the years.

#### References:

- Banks, T., 1990, *IBVS*, 3455  
Bakiş, H., Bakiş, V., Eker, Z., Demircan, O., 2016, *MNRAS*, **458**, 508 DOI  
Budding, E., 1989, *Space Science Reviews*, **50**, 205 DOI  
Collins, K. A., Collins, K. I., Pepper, J., et al., 2018, *ApJ*, **156**, 234 DOI  
Forbes, M., Budding, E., Priestley, J., 1988, *IBVS*, 3278  
Gaposchkin, S., 1953, *Annals of Harvard Coll. Obs.*, **113**, 67  
Kondo, Y., McCluskey, G.E. Jr., Parsons, S. B., 1985, *ApJ*, **295**, 580 DOI  
Nield, K. M., Priestley, J., Budding, E., 1986, *IBVS*, 2491  
Piotrowski, S. L., Ruciński, S. M., Semeniuk, I., 1974, *AcA*, **24**, 389  
Reed, P. A., 2011, *IBVS*, 5975  
Reed, P. A., 2012, *IAUS*, **282**, 325 DOI  
Roberts, A. W., 1894, *AJ*, **14**, 113 DOI  
Wittenmyer, R., Horner, J., Carter, B. D., et al., 2018, arXiv:1806.09282