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N49: THE FIRST ROBUST DISCOVERY OF RECOMBINING PLASMA IN AN EXTRA GALACTIC SUPERNOVA REMNANT

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ABSTRACT

Recent discoveries of recombining plasmas (RPs) in supernova remnants (SNRs) have dramatically changed our understanding of SNR evolution. To date, a dozen RP SNRs have been identified in the Galaxy. Here, we present *Suzaku* deep observations of four SNRs in the LMC, namely, N49, N49B, N23, and DEM L71, for accurate determination of their plasma states. Our uniform analysis reveals that of these SNRs, only N49 is in the recombining state, which is the first robust discovery of an RP from an extra-galactic SNR. Given that RPs have only been identified in core-collapse SNRs, our results strongly suggest a massive star origin for this SNR. On the other hand, no clear evidence of an RP is confirmed in N23, for which detection of recombination lines and continua was previously claimed. Comparing the physical properties of the RP SNRs identified so far, we find that all of them can be categorized into the “mixed-morphology” class, interacting with surrounding molecular clouds. This might be a key to determining the formation mechanisms of RPs.

Key words: ISM: abundances – ISM: individual objects (DEM L71, N23, N49, N49B) – ISM: supernova remnants – X-rays: ISM

1. INTRODUCTION

X-ray observations of supernova remnants (SNRs) allow us to make accurate measurements of plasma conditions and elemental abundances in the supernova (SN) ejecta, providing unique insights into the progenitor’s evolution and explosion as well as the dynamical evolution of the remnants themselves. The hot plasmas in X-ray-emitting SNRs are commonly in a state of non-equilibrium ionization (NEI) where the ionization degrees of heavy elements are inconsistent with those expected for a collisional ionization equilibrium (CIE) plasma with a certain electron temperature (e.g., Masai 1984). It was widely believed for a long time that SNR plasma is always in the low-ionization state, and slowly ionizing to achieve CIE (hereafter, ionizing plasma (IP)). In fact, recent X-ray observations with sensitive satellites, like *Suzaku* (Mitsuda et al. 2007), have confirmed that the immediate post-shock gas in young SNRs indeed consists of extremely low ionized atoms together with hot electrons (e.g., Uchida et al. 2013; Yamaguchi et al. 2014b).

However, earlier *ASCA* observations had suggested the presence of recombining plasma (RP; where the atoms are overionized compared to the observed electron temperature) in a couple of SNRs, IC 443 and W49B (Kawasaki et al. 2002, 2005). Conclusive evidence for the RPs was later revealed by *Suzaku* observations that discovered enhanced radiative recombination continua (RRCs) in the X-ray spectra of these remnants (Ozawa et al. 2009; Yamaguchi et al. 2009). These results were followed by observational studies of other SNRs, leading to a significant increase in the number of RP SNRs. To date, RPs have been discovered in a dozen SNRs, which means that the presence of RPs is no longer unusual or surprising.

The well-identified RP SNRs are all categorized into the so-called mixed-morphology (MM) class, defined by centrally

filled thermal X-ray emission with a synchrotron radio shell (Rho & Petre 1998). While more than 25% of the X-ray-detected Galactic SNRs are classified into this type (Jones et al. 1998), the physical process occurring to form such characteristic morphology is still unclear. Note that, theoretically, dense ambient materials play an important role in the dynamical evolutions of MM SNRs (e.g., White & Long 1991; Petruk 2001). Notably, another prediction is that the formation of RPs can also be explained by the interaction between the SN ejecta and the dense materials surrounding the progenitor (e.g., Itoh & Masai 1989). Therefore, the physical processes that create the MM and RP seem to be related to each other.

It is worth noting that most of the RP SNRs identified thus far were previously considered to have IP or nearly CIE plasma (e.g., W44; Uchida et al. 2012). This implies that there are still a number of SNRs for which the plasma state has been misclassified. Here, we present a uniform analysis of high-quality data of MM SNRs (as well as typical shell-like SNRs simultaneously observed) in the Large Magellanic Cloud (LMC) obtained using the X-ray Imaging Spectrometer (XIS; Koyama et al. 2007) on board *Suzaku*, in order to search for RPs in these SNRs. The LMC is particularly suitable for such a systematic study because of its low foreground extinction (Dickey & Lockman 1990) and its known distance of 50 kpc (Feast 1999); we use this distance throughout the paper.

In Section 2, we outline our observations and data reduction procedures. The results of the spectral analysis of four LMC SNRs, namely, N49, N23, N49B, and DEM L71, are presented in Section 3 where we will also show the new discovery of an RP from N49. We discuss our results in Section 4, and finally provide conclusions and future prospects in Section 5. The errors quoted in text, tables, and figures are at the 90% confidence level unless otherwise stated.

Table 1
Observation Logs

Target	Obs. ID	Obs. Date	(R.A., decl.) _{J2000}	Exposure
N23 & DEM L71	807008010	2012 Apr 4	(76.45, -67.96)	102 ks
N49 & N49B	807007010	2012 May 9	(81.50, -66.08)	185 ks

2. OBSERVATIONS AND DATA REDUCTION

We observed two regions in the LMC aimed at the MM SNRs N23 and N49 during the *Suzaku* Cycle 8 phase. Detailed information of the observations is summarized in Table 1. In addition to the targeted sources, the nearby SNRs DEM L71 and N49B are detected in each field of view (FOV) of the XIS. We also analyzed their spectra for a comparison of the plasma states among the targets. Both DEM L71 and N49B are classified as standard shell-like SNRs (Hughes et al. 2003; Park et al. 2003b) from which an RP has never been observed to date.

We analyzed data from three active CCDs: one is back-illuminated (XIS1) and the others are front-illuminated (FI; XIS0 and XIS3), although only the merged and averaged FI CCD spectra are shown in figures in the subsequent sections. We used HEASoft tools version 6.11 for the data reduction. The uncleaned data were reprocessed using the calibration database released in 2011 September, and screened with the standard event selection criteria for cleaned events. The resulting effective exposures were 102 and 185 ks for the observations of N23 (with DEM L71) and N49 (with N49B), respectively.

3. ANALYSIS AND RESULTS

In Figure 1, we show the 0.6–10.0 keV images of the two observations after subtraction of non X-ray background (NXB) generated using *xisnxbgen* (Tawa et al. 2008). Since the diameters of the SNRs ($\lesssim 1'.5$; e.g., N49B; Park et al. 2003b) are smaller than the angular resolution of the X-Ray Telescopes (the half-power diameter ranges from $1'.8$ to $2'.3$), our analysis focuses on spectra from the entire SNRs. We extract the source spectra from the regions enclosed by the circles shown in Figure 1. The background spectra are taken from the entire FOV excluding the source regions as well as the CCD corners illuminated by the ^{55}Fe calibration sources.

Figure 2 compares the background-subtracted spectra of the four SNRs, showing that N49 is the brightest among them (with a count rate about ~ 10 times higher than the others). Due to the good energy resolution of the XIS, we clearly detect K-shell emission lines of Ar, Ca, and Fe from this remnant for the first time. The spectrum of N49 exhibits strong Ly α lines of Mg, Si, and S, whereas there is little or no signal of these lines from the other remnants. This indicates that the heavy elements in N49 are more highly ionized than those in the others. Since our main goal is to identify an RP SNR(s), we first perform detailed spectral analysis of N49 (Section 3.1), followed by an analysis of the other remnants (Sections 3.2–3.4).

We use the SPEX software version 2.04.01 (Kaastra et al. 1996) for the spectral fitting by taking into account the detector and telescope responses (the so-called “redistribution matrix” and “ancillary response,” respectively) generated by *xisrmfgen* and *xissimarfgen* (Ishisaki et al. 2007). The data around the neutral Si K-shell edge (1.77–1.83 keV) are ignored because of the poor accuracy in the response function

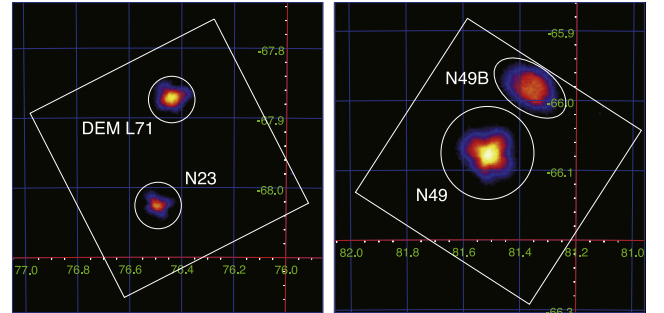


Figure 1. Vignetting-corrected XIS FI images of the N23 (left) and N49 (right) regions in the 0.6–10.0 keV band with an equatorial coordinate grid. DEM L71 and N49B are also detected in the FOV of the XISs indicated by the solid squares.

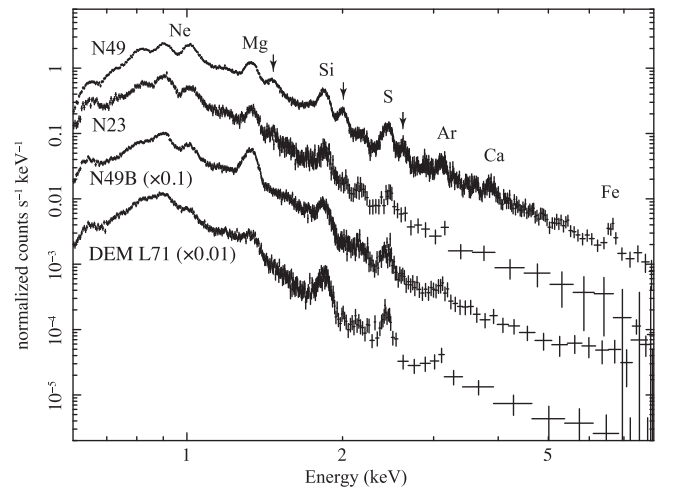


Figure 2. XIS FI spectra of N49, N23, N49B, and DEM L71 where the background spectra are subtracted. The spectra of N49B and DEM L71 are multiplied by factors of 0.1 and 0.01, respectively. The arrows represent the line centroids of Mg-Ly α , Si-Ly α , and S-Ly α .

at these energies.⁵ During the analysis, we separately set absorption column densities in the Milky Way ($N_{\text{H(MW)}}$) and the LMC ($N_{\text{H(LMC)}}$). The former value is fixed to $6 \times 10^{20} \text{ cm}^{-2}$ (Dickey & Lockman 1990). The elemental abundances of the latter component are fixed to the average values of the LMC (~ 0.3 solar; Russell & Dopita 1992).

3.1. N49

N49 is one of the brightest SNRs in the LMC at the various wavelengths observed (e.g., Long et al. 1981; Dickel & Milne 1998). The radio continuum and H α emission exhibit a clear shell structure (Vancura et al. 1992; Dickel et al. 1995), whereas the X-ray image shows an irregular morphology that is brightest inside the radio shell in the southwest region (Park et al. 2012), as is a typical characteristic of an MM SNR. N49

⁵ <http://heasarc.nasa.gov/docs/Suzaku/analysis/sical.html>

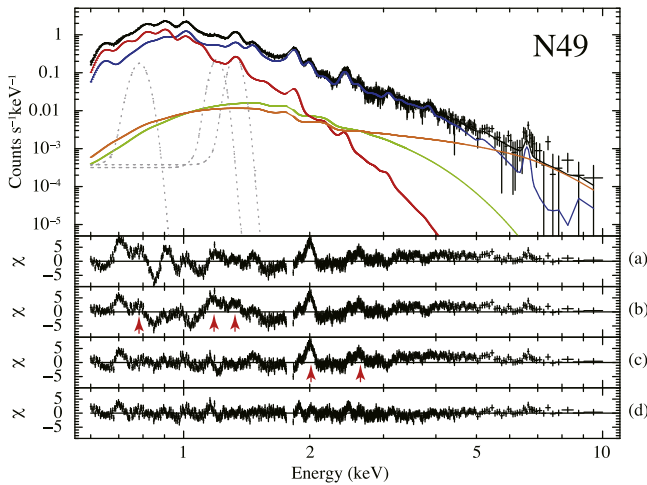


Figure 3. XIS FI spectrum of N49 (top panel). The best-fit model is overlaid with the solid black line. The solid blue and red lines represent the best-fit high- kT_e (RP) and the low- kT_e (IP) components, respectively. The orange and green lines show the power-law and the blackbody components for SGR 0526–66, respectively. The dotted lines show Gaussians for the Fe L emission missing from the plasma code we used. The residuals are shown in panel (d). Panels (a), (b), and (c) represent the residuals from the models of single IP, two IP, and two IP plus Gaussians, respectively.

spatially overlaps with the soft gamma-ray repeater (SGR) 0526–66 (Cline et al. 1982), although the physical association between the SGR and SNR is under debate (Gaensler et al. 2001; Badenes et al. 2009). Using *Chandra* data, Park et al. (2012) revealed that the X-ray spectrum of SGR 0526–66 is best reproduced by a model consisting of a blackbody (BB) with $kT_{\text{BB}} = 0.44 \pm 0.02$ keV and a power law (PL) with $\Gamma = 2.50^{+0.11}_{-0.12}$. Unfortunately, our *Suzaku* data do not spatially resolve the SGR from the SNR. We thus add the best-fit model by Park et al. (2012) to the model spectrum of N49, but allow its spectral parameters to vary within the reported statistical uncertainty. Since Park et al. (2012) found significant long-term variability in the X-ray flux of the SGR, we also allow the fluxes of both BB and PL components to vary freely.

First, we apply a single-component IP model where the initial ionization balance is dominated by the neutral state for all heavy elements. The free parameters are the electron temperature (kT_e), ionization timescale ($n_e t$, where n_e and t are the electron number density and the elapsed time since the gas was shock heated), emission measure ($\text{EM} = \int n_e n_H V$), and abundances of O, Ne, Mg, Si, S, Ar, Ca, and Fe relative to the solar values of Anders & Grevesse (1989). The Ni abundance is linked to the value of Fe, while the abundances of the other elements are fixed to the LMC averages of Russell & Dopita (1992). This model yields a best-fit electron temperature of ~ 0.6 keV but fails to reproduce the overall spectrum with an unacceptable $\chi^2/\text{degrees of freedom}$ (dof) of 4936/820. The residual is shown in panel (a) of Figure 3. The particularly large disagreement between the data and model is seen below ~ 1 keV, implying the presence of another plasma component with a lower temperature. Assuming that this component originates from a swept-up interstellar medium (ISM) that is dominant at the outermost region of the remnant (Park et al. 2003a, 2012), we added another IP component with the abundances fixed to the LMC mean values. Although this two-component model gives a slightly better fit ($\chi^2/\text{dof} = 4297/817$), some line-like residuals still remain, as shown in panel

(b) of Figure 3. The residual at ~ 0.8 keV, also observed in our previous studies of other objects (e.g., Nakashima et al. 2013; Uchida et al. 2013), is likely due to the well-known uncertainty in the emissivity ratio between the Fe XVIII L-shell emission of $3s \rightarrow 2p$ and $3d \rightarrow 2p$ (Gu et al. 2007). On the other hand, the residuals at ~ 1.2 and ~ 1.3 keV are caused by L-shell emission from high quantum numbers ($n > 5$) missing from the plasma code SPEX (Brickhouse et al. 2000). We compensate for these lines by adding three Gaussians at these energies and obtain the improved fit ($\chi^2/\text{dof} = 2366/815$) with the best-fit kT_e of ~ 0.6 keV and $n_e t$ of $\sim 10^{13} \text{ cm}^{-3} \text{ s}$ for the high- kT_e component. The obtained ionization timescale indicates that this IP model actually represents a CIE plasma at the given temperature.

The fit left the largest residuals at 2.0 and 2.6 keV (see panel (c) of Figure 3), which correspond to the centroid energies of the Si Ly α and S Ly α emissions, respectively. This suggests that the average charge of these elements is higher than that expected in a 0.6 keV CIE plasma. Furthermore, a hump-like feature is found in the residual around 2.7 keV, indicating an enhancement of the RRC of Si, as was observed in the other RP SNRs (e.g., IC 443; Yamaguchi et al. 2009). Therefore, we introduce an RP model (using an “NEIJ” model in SPEX⁶) for the high- kT_e component. In addition to the parameters given in the IP model, the RP model has another free parameter kT_{init} higher than kT_e , which describes a recombining history by $n_e t$ in the kT_e plasma starting from the initial ionization temperature of kT_{init} . This model dramatically improves the fit ($\chi^2/\text{dof} = 1270/813$), removing the large residuals above 2 keV as shown in the panel (d) of Figure 3. The best-fit model components and parameters are given in Figure 3 and Table 2, respectively.

Interestingly, the RP model reproduces the Fe K α emission at ~ 6.6 keV as well, although free electrons in the 0.6 keV plasma are not energetic enough to excite K-shell electrons of Fe. This indicates that the observed Fe K α emission predominantly originates from cascade processes after radiative recombination into the excited levels of He-like ions (i.e., $\text{Fe}^{25+} + e^- \rightarrow \text{Fe}^{24+*}$). We emphasize that this is why kT_{init} is constrained to the very high value ($\gtrsim 10$ keV) so that a significant fraction of the Fe ions remains at the H-like state in the current plasma. It should also be noted that the observed centroid energy 6629^{+31}_{-26} eV is consistent with the value expected from the best-fit RP model (6658 eV) but significantly lower than that for a typical CIE plasma at $kT_e = 5$ –10 keV (~ 6680 eV). This is another piece of evidence for the RP; the forbidden and intercombination lines are likely to be enhanced by the recombination processes.

3.2. N23

In the previous work using *XMM-Newton* MOS and RGS data, Broersen et al. (2011) claimed that the X-ray spectrum of N23 can be best reproduced by a two-component NEI model of which one of the components was an RP with $kT_{\text{init}} = 3.0$ keV and $kT_e = 0.18$ keV. We first fit the XIS spectrum fixing their best-fit parameters but fail to reproduce the XIS spectrum ($\chi^2/\text{dof} = 455/236$), leaving large residuals around the Ne-Ly α emission and the Fe XVII recombination edge at ~ 1.25 keV. The fit is not improved even if we allow the abundances to vary. Therefore, we thaw the temperatures and ionization parameters (but restricting kT_{init} to be higher than kT_e), obtaining a

⁶ See <http://www.sron.nl/files/HEA/SPEX/manuals/manual.pdf> for details.

Table 2
Best-fit Parameters

Component	Parameter	N49	N23	N49B	DEM L71
Absorption	$N_{\text{H(LMC)}} (\times 10^{21} \text{ cm}^{-2})$	$3.51_{-0.04}^{+0.01}$	2.13 ± 0.05	$2.30_{-0.06}^{+0.01}$	$3.70_{-0.50}^{+0.30}$
	$N_{\text{H(MW)}} (\times 10^{21} \text{ cm}^{-2})$	0.6 (fixed)	0.6 (fixed)	0.6 (fixed)	0.6 (fixed)
ISM	kT_e (keV)	0.30 ± 0.01	0.17 ± 0.01	0.21 ± 0.01	0.16 ± 0.01
	$n_e t (\times 10^{11} \text{ cm}^{-3} \text{ s})$	$3.49_{-0.26}^{+0.09}$	> 10	$1.23_{-0.07}^{+0.09}$	> 10
	EM ($\times 10^{58} \text{ cm}^{-3}$)	$3.54_{-0.01}^{+0.03}$	17.6 ± 0.3	6.0 ± 0.9	19_{-4}^{+5}
Ejecta	kT_e (keV)	0.62 ± 0.01	0.61 ± 0.01	0.79 ± 0.01	0.69 ± 0.01
	kT_{init} (keV)	11 ± 1	0.01 (fixed)	0.01 (fixed)	0.01 (fixed)
	$n_e t (\times 10^{11} \text{ cm}^{-3} \text{ s})$	$7.00_{-0.02}^{+0.44}$	$5.08_{-0.37}^{+0.42}$	0.84 ± 0.01	1.77 ± 0.07
	O	$1.52_{-0.03}^{+0.04}$	3.5 ± 0.2	$0.63_{-0.02}^{+0.03}$	$0.30_{-0.04}^{+0.05}$
	Ne	$0.98_{-0.02}^{+0.03}$	1.25 ± 0.07	0.45 ± 0.01	0.25 ± 0.02
	Mg	0.82 ± 0.02	1.16 ± 0.07	1.24 ± 0.02	0.62 ± 0.03
	Si	$1.32_{-0.02}^{+0.05}$	1.2 ± 0.4	0.34 ± 0.02	0.48 ± 0.04
	S	$1.46_{-0.11}^{+0.03}$	1.25 ± 0.08	0.50 ± 0.08	1.0 ± 0.2
	Ar	$1.3_{-0.5}^{+0.1}$	< 3.1	0.7 ± 0.6	$1.5_{-0.3}^{+3.1}$
	Ca	$1.5_{-0.9}^{+0.2}$	(=Ar)	(=Ar)	(=Ar)
	Fe (=Ni)	0.24 ± 0.01	0.46 ± 0.02	0.35 ± 0.01	0.62 ± 0.01
	EM ($\times 10^{58} \text{ cm}^{-3}$)	1.73 ± 0.01	0.62 ± 0.01	0.79 ± 0.01	1.36 ± 0.01
		χ^2/dof	1270/813	218/234	659/423

significantly better fit with $\chi^2/\text{dof} = 224/234$ and $n_e t > 10^{12} \text{ cm}^{-3} \text{ s}$ for both components. This high $n_e t$ value indicates that the plasmas are in nearly CIE independent of the initial ionization population (otherwise ionizing). In fact, we are not able to constrain kT_{init} with this model. Finally, we apply an IP model consisting of an LMC-abundance component (for ISM) and a free-abundance component (for ejecta), and obtain a slightly better fit ($\chi^2/\text{dof} = 218/234$) than the RP (or CIE) model. The best-fit parameters and models are given in Table 2 and Figure 4, respectively. During the analysis, we fix the spectral shape and flux of the faint X-ray compact source CXOU J050552.3–680141 associated with N23 to the reported values of Hayato et al. (2006), although its contribution is negligible at $\lesssim 3 \text{ keV}$ (see Figure 4).

3.3. N49B

N49B is a typical shell-like SNR containing no X-ray point-like source within the SNR shell (Park et al. 2003b). We fit the XIS spectrum with the same two-component IP model applied to N49, obtaining the best-fit results given in Table 2. The $n_e t$ value for the ejecta component ($\sim 8 \times 10^{10} \text{ cm}^{-3} \text{ s}$) is the lowest among the SNRs studied in this work. We reveal that the abundance of Mg is significantly higher than those of the other elements (e.g., $\text{Mg/O} \sim 2.0$, $\text{Mg/Ne} \sim 2.8$, $\text{Mg/Si} \sim 3.6$). Park et al. (2003b) previously found “Mg-rich” ejecta near the center of the remnant. Our result shows that the Mg enrichment is confirmed even in the integrated spectrum from the entire SNR. We find no evidence for an RP; if $kT_{\text{init}} > kT_e$ is assumed, then the model clearly fails to reproduce the spectrum with $\chi^2/\text{dof} = 1582/426$. Therefore, we conclude that this SNR is dominated by IPs.

3.4. DEM L71

DEM L71 is a middle-aged, shell-like SNR where a double-shock morphology was observed by *Chandra* observations (Hughes et al. 2003), suggesting the presence of both reverse-

shocked ejecta and swept-up ISM. Hughes et al. (2003) also revealed that the X-ray spectra from the interior regions are dominated by strong Fe L emission. Similar to the spectral analysis of N49B, we fit the XIS spectrum of DEM L71 with a two-component IP model consisting of a low- kT_e ISM and a high- kT_e ejecta. The fit is acceptable with χ^2/dof of 291/229 as is given in Table 2. The Fe abundance is highest among the four SNRs, consistent with previous studies (e.g., Hughes et al. 1998, 2003) where the Type Ia origin of this SNR was suggested. If we assume $kT_{\text{init}} > kT_e$, then an unacceptable fit ($\chi^2/\text{dof} = 399/228$) is obtained, ruling out an RP scenario for this remnant.

4. DISCUSSION

4.1. Plasma Conditions in the Observed SNRs

We have systematically analyzed X-ray spectra from the four LMC SNRs (N49, N23, N49B, and DEM L71) and revealed robust evidence for overionization in N49—this is the first discovery of an RP from an extragalactic SNR. The other SNRs, including N23 from which the presence of an RP was previously reported (Broersen et al. 2011), can be fairly well characterized by an IP or a nearly CIE plasma.

The previous claim of the recombining state in N23 was based on the enhanced G -ratio of the O VII lines measured using *XMM-Newton* RGS. Although Broersen et al. (2011) claimed that the broadband spectrum of *XMM-Newton* MOS can also be reproduced by an RP model, their best-fit model is clearly ruled out by our analysis of the XIS spectrum with better photon statistics. Since our analysis excluded the O VII emission, it is still possible that the swept-up ISM (low- kT_e component) is partially recombining. We note, however, that there are several other processes that can enhance the G -ratio: resonance scattering and charge exchange. Notably, a similar G -ratio enhancement in the O VII lines was observed in DEM L71, and was indeed interpreted to be a consequence of resonance scattering (van der Heyden et al. 2003). Since the charge

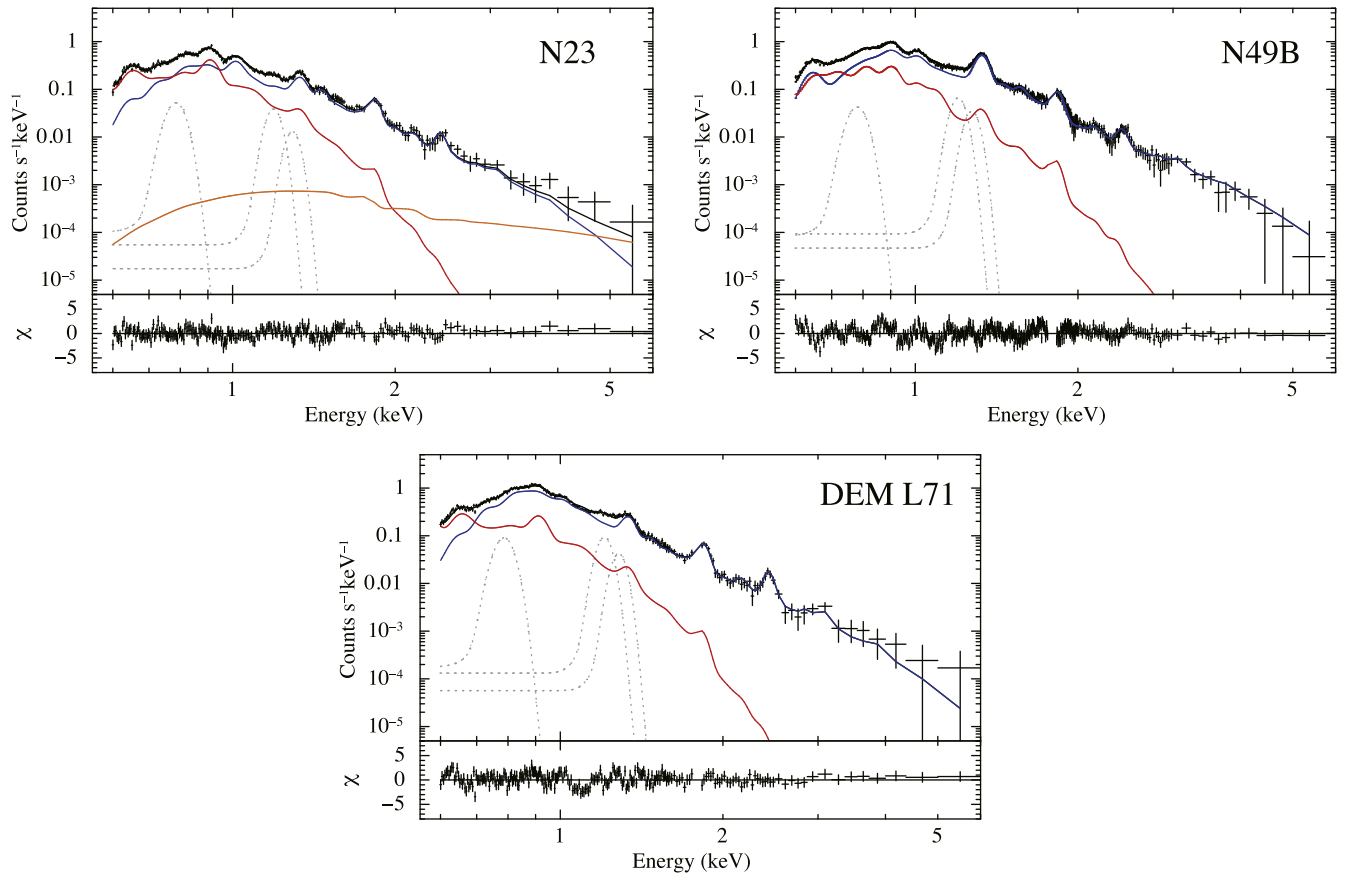


Figure 4. XIS FI spectra of N23 (upper left), N49B (upper right), and DEM L71 (bottom). The best-fit models are overlaid with the black solid lines. The solid blue and red lines represent the best-fit high- kT_e and the low- kT_e components, respectively. The dotted lines show Gaussians for the Fe L emission missing from the plasma code we used. The orange line in the model of N23 shows the power-law component for CXOU J050552.3–680141 (see the text).

Table 3
Summary of the Four LMC SNRs Studied in This Work

SNR	RP	Age (10^4 years)	SN Type	Morphology	Cloud Interaction	References
N49	Yes	0.4–0.5	Core-collapse	MM	Yes (CO and H α)	This work, 1, 2, 3
N23	No	0.4	Core-collapse	MM	No	2, 4
N49B	No	1.09	Core-collapse	Shell	No	5
DEM L71	No	0.4–0.5	Ia	Shell	No	6, 7, 8

References. (1) Park et al. (2012), (2) Banas et al. (1997), (3) Melnik & Copetti (2013), (4) Broersen et al. (2011), (5) Park et al. (2003b), (6) Ghavamian et al. (2003), (7) Hughes et al. (1998), (8) Hughes et al. (2003).

exchange process seems to work in some evolved SNRs (e.g., Puppis A; Katsuda et al. 2012), this process may also be responsible for the high G -ratio observed in N23.

4.2. Origin of the RP

We summarize the characteristics of the four SNRs studied in this work in Table 3. Only N49 is interacting with dense clouds identified by the CO (Banas et al. 1997) and H α (Melnik & Copetti 2013) observations. Given that similar cloud interaction is observed in most SNRs from which the presence of an RP has been confirmed (e.g., W49B; Ozawa et al. 2009), thermal conduction into surrounding clouds might play an important role in forming the RP (see Zhou et al. 2011, for a discussion). In this scenario, there should be a spatial correlation between the electron temperature of the SNR plasma and the cloud density. Future deep observations of N49

with better spatial resolution are necessary to assess this possibility.

An alternative is that the adiabatic expansion of the SNR caused a rapid cooling of electrons and the resulting recombining state of the plasma. In this scenario, the SN ejecta should interact with dense circumstellar matter (CSM) in the early phase of the SNR evolution, and the highly ionized ejecta and CSM expand drastically after the SNR blast wave breaks out to the low-density ISM (Itoh & Masai 1989). Therefore, the ionization timescale ($n_e t$) in this scenario is characterized by the elapsed time and the density evolution history since the break out took place. To evaluate the possibility of this scenario, we estimate the electron density in the present RP of N49 using the derived EM ($=n_e n_H V$). Since the plasma density is known to be highly inhomogeneous in this remnant (Park et al. 2012), we analyze archival *Chandra* data of N49 to investigate the surface brightness profile. We

Table 4
List of the RP SNRs in Order of Increasing Electron Temperature

SNR	Age (10^4 years)	kT_{init} (keV)	kT_e (keV)	$n_e t$ ($10^{11} \text{ cm}^{-3} \text{ s}$)	Cloud Interaction	Compact Object	References
G359.1–0.5	>1	$0.87^{+0.15}_{-0.11}$	0.29 ± 0.02	<4.42	Yes	...	1, 2, 24
G346.6–0.2	...	5 (fixed)	$0.30^{+0.03}_{-0.01}$	$4.8^{+0.1}_{-0.4}$	Yes	...	3, 25
W28	~ 4	3 (fixed)	$0.40^{+0.02}_{-0.03}$	6.31	Yes	PSR 1758–23	4, 5, 6, 24
W44	~ 2	$1.07^{+0.08}_{-0.06}$	0.48 ± 0.02	6.76 ± 0.5	Yes	PSR B1853+01	7, 8, 9, 24
CTB 37 A	~ 1	5 (fixed)	$0.49^{+0.09}_{-0.06}$	13^{+3}_{-1}	Yes	CXOU J171419.8–383023	10, 11, 12, 24
3C391	...	$1.8^{+1.6}_{-0.6}$	0.495 ± 0.015	$14.0^{+1.5}_{-2.2}$	Yes	...	13, 24
MSH 11–61 A	1–2	5 (fixed)	$0.513^{+0.004}_{-0.003}$	12.2 ± 0.4	Possible	IGR J11014–6103	14, 15, 16, 26
N49	0.4–0.5	11 ± 1	0.62 ± 0.01	$7.00^{+0.44}_{-0.02}$	Yes	SGR 0526–66	This work, 17
IC 443	~ 0.4	10 (fixed)	0.65 ± 0.04	9.8 ± 1.1	Yes	XMMU J061804.3+222732	18, 19, 20, 21, 24
W49B	~ 0.4	...	$1.52^{+0.01}_{-0.02}$...	Yes	...	22, 23, 27

Note. The plasma parameters for IC 443 and W49B were determined only from a high-energy band around Fe-K lines (e.g., Ozawa et al. 2009; Ohnishi et al. 2014), while those for the other SNRs were determined from intermediate-mass elements.

References. (1) Aharonian et al. (2008b), (2) Ohnishi et al. (2011), (3) Yamauchi et al. (2013), (4) Rho & Borkowski (2002), (5) Sawada & Koyama (2012), (6) “runaway” Pulsar; Frail et al. (1993), (7) Cox et al. (1999), (8) Uchida et al. (2012), (9) Pulsar plus Pulsar Wind Nebula; Petre et al. (2002), (10) Yamauchi et al. (2014), (11) Wolszczan et al. (1991), (12) Pulsar Wind Nebula candidate; Aharonian et al. (2008a), (13) Sato et al. (2014), (14) Slane et al. (2002), (15) Kamitsukasa et al. (2015), (16) Pulsar plus Pulsar Wind Nebula; Pavan et al. (2014), (17) Park et al. (2012), (18) Lee et al. (2008), (19) Yamaguchi et al. (2009), (20) Ohnishi et al. (2014), (21) Neutronstar plus Synchrotron Nebula; Olbert et al. (2001), (22) Hwang et al. (2000), (23) Ozawa et al. (2009), (24) Hewitt et al. (2008), (25) Koralesky et al. (1998), (26) Filipović et al. (2005), (27) Keohane et al. (2007).

find that $\sim 50\%$ of the photon flux in the Si band (1.7–2.1 keV) is coming from the brightest southeast region (approximately a $0.4 \times 0.25 \times 0.25$ ellipsoid), and the other $\sim 50\%$ is from the remaining faint region in the entire SNR (a sphere with a radius of 0.6). Therefore, the volumes and the electron densities are estimated to be 9.6×10^{57} and 1.0 cm^{-3} for the bright region, and 7.3×10^{58} and 0.38 cm^{-3} for the faint region at a distance of 50 kpc. If we simply divide the best-fit $n_e t$ in the ejecta component ($7.0 \times 10^{11} \text{ cm}^{-3} \text{ s}$) by the derived electron density, then we obtain plasma ages of 22–58 kyr, more than a few times larger than the dynamical age of N49 (~ 4800 years; Park et al. 2012). This is not surprising because the plasma must have had a higher electron density in the past. Yamaguchi et al. (2012) also obtained a similar result from the *Suzaku* observation of IC 443 where the plasma age estimate (~ 11 kyr) is significantly higher than the SNR age of ~ 4000 years (Troja et al. 2008).

In Table 4, we summarize the physical properties of the RP SNRs identified so far. While the estimated ages roughly correlate with kT_e , the values of $n_e t$ show no clear trend. This may be partly due to the technical reason that the parameter $n_e t$ is coupled to that of kT_{init} in the spectral fitting. Nevertheless, Table 4 shows that the $n_e t$ values exceed $\sim 10^{11} \text{ cm}^{-3} \text{ s}$ in all SNRs, implying that the overionization did *not* start from the last few hundred years, but from the early epoch in the SNR. This fact likely favors the adiabatic cooling scenario for most RP SNRs.

4.3. Progenitor of N49 and SGR 0526–66

The SN type of N49 has been controversial. Park et al. (2012) suggested a Type Ia origin based on the Si/S abundance ratio in the ejecta measured by the *Chandra* data. On the other hand, the environment of N49, i.e., a nearby OB association (Chu & Kennicutt 1988) and young stellar clusters (Klose et al. 2004), are more common in core-collapse SNRs. A recent systematic study of Fe K emission in young and middle-aged SNRs also agrees on a core-collapse origin (Yamaguchi et al. 2014a). Given that an RP has never been observed in a Type Ia

SNR, our results strongly support a core-collapse scenario for this SNR. It should also be noted that the formation of an RP requires dense CSM, which is only made by a massive progenitor with a significant mass loss rate (Moriya 2012; Shimizu et al. 2012). The total mass of the RP component is roughly estimated to be $\sim 26 M_\odot$ using the density and emitting volume derived in Section 4.2, which is in marginal agreement with the progenitor mass of 12.5–21.5 M_\odot constrained from the local star formation history around N49 (Badenes et al. 2009).

Since a core-collapse origin is strongly suggested for N49, the compact object SGR 0526–66, which is a known magnetar candidate (e.g., Park et al. 2012), might be associated with the SNR. The large distance between SGR 0526–66 and the geometric center of N49 requires a high kick velocity ($\sim 1100 \text{ km s}^{-1}$) for the SGR at an SNR age of ~ 4800 years, and hence Park et al. (2012) argued against the physical association between these objects. However, the theoretical prediction is that a magnetorotational core-collapse SN explodes asymmetrically with bipolar jets and produces a magnetar with a large kick velocity up to $\sim 1000 \text{ km s}^{-1}$ (Sawai et al. 2008). Moreover, there are several compact objects located extremely off-center from the associated RP SNRs. For instance, IGR J11014–6103 is located outside the shell of MSH 11–61 A. Pavan et al. (2014) indicated that this neutron star is escaping from MSH 11–61 A at a velocity exceeding 1000 km s^{-1} . Also, PSR 1758–23 is known as a “runaway pulsar” which might be physically associated with W28 (Frail et al. 1993). Two other compact objects, CXOU J171419.8–383023 in CTB 37 A (Aharonian et al. 2008a) and XMMU J061804.3+222732 in IC 443 (Olbert et al. 2001), are also found far from the geometric center of their host SNRs. Taking these facts into consideration, we propose that SGR 0526–66 is the plausible candidate for the compact remnant of N49.

5. CONCLUSIONS

We have shown clear evidence for an RP in the SNR N49, which is the first report of the robust detection of an RP from an

extra-galactic SNR achieved by the high sensitivity and good energy resolution of the XIS on board *Suzaku*. The other SNRs, including N23 where the presence of an RP was previously claimed (Broersen et al. 2011), can be well characterized by an IP or nearly CIE plasma. Future observation of N49 using high-resolution X-ray spectrometers, like *ASTRO-H* (Takahashi et al. 2014), is highly encouraged; our *Suzaku* result predicts the detection of the Fe RRC as well as enhanced forbidden lines due to recombination processes (see Section 3.1). We also emphasize the importance of a systematic study of RP SNRs, which will help us to understand the physical processes responsible for the formation of the RPs in more detail.

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