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THE EFFECT OF MINOR AND MAJOR MERGERS ON THE EVOLUTION OF LOW EXCITATION RADIO GALAXIES

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ABSTRACT

We use deep, $\mu_r \lesssim 28$ mag arcsec⁻², r -band imaging from the Dark Energy Camera Legacy Survey (DECaLS) to search for past, or ongoing, merger activity in a sample of 282 Low Excitation Radio Galaxies (LERGs) at $z < 0.07$. Our principle aim is to assess the the role of mergers in the evolution of LERGs. Exploiting the imaging depth, we classify tidal remnants around galaxies as both minor and major morphological disturbances for our LERG sample and 1,622 control galaxies matched in redshift, stellar mass, and environment. In groups and in the field, the LERG minor merger fraction is consistent with the control population. In galaxy clusters, $8.8 \pm 2.9\%$ of LERGs show evidence of recent minor mergers in contrast to $23.0 \pm 2.0\%$ of controls. This $\sim 4\sigma$ deficit of minor mergers in cluster LERGs suggests these events may inhibit this type of nuclear activity for galaxies within the cluster environment. We observe a $> 4\sigma$ excess of major mergers in the LERGs with $M_* \lesssim 10^{11} M_\odot$, with $10 \pm 1.5\%$ of these AGN involved in such large-scale interactions compared to $3.2 \pm 0.4\%$ of control galaxies. This excess of major mergers in LERGs decreases with increasing stellar mass, vanishing by $M_* > 10^{11.3} M_\odot$. These observations show that minor mergers do not fuel LERGs, and are consistent with typical LERGs being powered by accretion of matter from their halo. Where LERGs are associated with major mergers, these objects may evolve into more efficiently accreting active galactic nuclei as the merger progresses and more gas falls on to the central engine.

Keywords: galaxies: active — galaxies: interactions — galaxies: evolution — galaxies: nuclei

1. INTRODUCTION

The accretion of matter onto the central supermassive black hole within galaxies powers active galactic nuclei

(AGN, e.g., [Salpeter 1964](#); [Kaviraj et al. 2017](#)). When the accretion rate of matter, \dot{M} , is greater than $\sim 1\%$ of the Eddington accretion rate, \dot{M}_{Edd} , the accretion mode can be described as radiatively efficient (e.g., [Jackson & Rawlings 1997](#); [Hardcastle et al. 2007](#)). Such radiatively efficient accretion modes allow for the formation of an optically thick accretion disk that radiates high energy photons (i.e. ultraviolet with upscattering to x-ray, e.g.,

Baum et al. 1995; Heckman et al. 2005). These high energy photons act to ionise the local interstellar medium, producing the narrow excitation lines seen in, for example, optically selected AGN (e.g., Buttiglione et al. 2010).

In addition to the accretion of matter on to the nuclear black hole, in some AGN energy can be further extracted from the central engine by the production of a relativistic jet (Blandford & Znajek 1977). Non-thermal radiation from the jet, and its interaction with the surrounding medium, produces relatively bright radio emission in such AGN. The bulk of these radio-loud AGN (RLAGN) at low redshift show no evidence of high-excitation lines resulting from an optically thick nuclear accretion mechanism (Best & Heckman 2012). These Low Excitation Radio Galaxies (LERGs) are thus thought to not possess the same type of accretion disk associated with radiatively efficient accretion modes, and instead are fuelled by an optically thin, advection dominated accretion flow (Fabian & Rees 1995; Narayan & Yi 1995). In order to explain this radiatively inefficient accretion whilst still producing the AGN jet, very low Eddington scaled accretion rates, i.e. $\dot{M} \ll 0.01\dot{M}_{\text{Edd}}$ are invoked (e.g., Baum et al. 1992, 1995; Tadhunter et al. 1998; Hardcastle et al. 2006; Allen et al. 2006; Evans et al. 2011; Mingo et al. 2014).

The weakly accreting nature of LERGs is indicative of a poorer fuel supply than is available to more efficiently accreting AGN (e.g., Best & Heckman 2012; Ellison et al. 2015). Furthermore, LERGs are usually passive in terms of star formation, and associated with red, massive, early-type galaxies (e.g., Heckman et al. 1986; Best et al. 2005a; Kauffmann et al. 2008; Lofthouse et al. 2018). This deficiency of star formation adds further weight to the argument that LERGs lack a ready supply of cold gas. The environments that host LERGs are frequently observed to be over-dense, with LERGs often being brightest cluster galaxies (e.g., Hill & Lilly 1991; Zirbel 1997; Best et al. 2007; Ramos Almeida et al. 2013; Ching et al. 2017). In such dense environments, the hot intra-cluster medium (ICM) acts to inhibit the accretion of cold gas by galaxies (Davies et al. 2017).

Limited supplies of cold gas can be accreted onto galaxies in cluster cores via cooling flows (e.g., O’Dea et al. 1994, 2008; Edge 2001; Pipino et al. 2009; Donahue et al. 2011) and hence provide a potential fuel supply for an AGN. Such cooling flows may be enhanced by thermal instabilities in a dynamic ICM, a process known as chaotic cold accretion (CCA, Gaspari et al. 2013, 2017). Additionally, AGN driven outflows may cool as they expand, allowing gas to fall back on to, and drip-feed the central engine (e.g., Tremblay et al. 2016; Voit et al.

2017; Tremblay et al. 2018). Finally, galactic mergers present an obvious mechanism with which to introduce a cold gas fuel reservoir to power the AGN (Sanders et al. 1988; Weston et al. 2017), although evidence for this mechanism is mixed (e.g., Scott & Kaviraj 2014; Villforth et al. 2017). Given the expectation for a limited fuel supply in LERGs, then if mergers are involved in their triggering they should either be gas-poor, or else indirect triggers rather than a direct fuel supply. In this scenario, minor mergers present an attractive trigger mechanism for LERGs (e.g., Kaviraj 2014a; Pace & Salim 2014; Ellison et al. 2015; Martin et al. 2018), and might be expected given the excess of satellites observed around LERGs (Pace & Salim 2014). Such low-impact galactic collisions would provide a restricted gas supply that may fall short of initiating the radiatively efficient accretion modes associated with High Excitation Radio Galaxies (HERGs) and AGN selected from non-radio bands.

In this work we aim to test this last hypothesis that mergers, and in particular minor mergers, play a role in the evolution of LERGs. The low mass ratio involved in a minor merger ($\lesssim 1:4$, Lotz et al. 2010) results in a limited impact on the morphology of the primary, or recipient, galaxy in the merger. The morphology of the secondary (hereafter donor) galaxy is totally disrupted as it is absorbed by the recipient galaxy. In combination these effects can make detecting minor mergers problematic in the relatively shallow imaging obtained by typical wide-field galaxy surveys (Kaviraj 2010, 2014b). Instead, observational evidence for minor mergers presents as subtle low surface brightness (LSB) tidal features (e.g., tails, streams and halo shells), the results of stellar material stripped from the donor galaxy during its infall onto the recipient galaxy (e.g., Ramos Almeida et al. 2012; Kaviraj 2014a).

Detecting LSB features requires deeper imaging than is necessary for normal morphological studies of massive galaxies. A new range of wide-deep imaging surveys such as Stripe 82 of the Sloan Digital Sky Survey (SDSS, York et al. 2000; Fliri & Trujillo 2016), the Kilo Degree Survey (KiDS, de Jong et al. 2013), and the Dark Energy Camera Legacy Survey (DECaLS, Blum et al. 2016), are paving the way for large studies of LSB structures such as tidal features. One would expect that, should mergers be involved in the evolution of a galaxy in to a LERG, then an excess of merger signatures would be observed in LERGs compared to a control sample of galaxies. Exploiting the latest in deep imaging surveys allows this approach to be extended to minor mergers by comparing the rates of LSB morphological disruption. This method was used by Kaviraj (2014b) to demonstrate the signif-

icant role of minor mergers in fuelling star formation at low redshift, and it is this procedure we employ in this paper.

The layout of this paper is as follows. In Section 2, we detail the datasets used in this paper and describe the LERG and control sample selections. Section 3 describes the process of classifying the DECaLS images. We state our results in Section 4 and discuss the implications of these observations in Section 5. Section 6 is a summary of this work. Throughout this work we assume a standard flat Λ CDM cosmology with $h = 0.7$, $H_0 = 100h \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, $\Omega_\Lambda = 0.7$.

2. OBSERVATIONAL DATASETS AND SAMPLE SELECTION

To assess the role of mergers in the evolution of LERGs requires optical spectra and imaging, as well as radio observations. Optical spectroscopic data is obtained from SDSS data release 7 (DR7, Abazajian et al. 2009), containing $\sim 10^6$ spectra of galaxies and quasars across $\sim 8,000 \text{ deg}^2$ of the sky. Optical imaging is obtained from data release 5 (DR 5) of DECaLS, which broadly covers the region $-22^\circ < \delta < +34^\circ$, $|b| > 18^\circ$ at $r \lesssim 24 \text{ mag}$. In the $\sim 5,000 \text{ deg}^2$ where the survey footprints overlap, DECaLS provides imaging that is approximately 2 mags deeper than the standard SDSS optical imaging. Beyond point source detection, it is the enhanced ability of DECaLS over SDSS to detect low surface brightnesses that is important in this work. In comparison to the standard depth SDSS imaging which can detect surface brightnesses of $\mu_r \sim 25 \text{ mag arcsec}^{-2}$ (Driver et al. 2016), DECaLS observes surface brightnesses of $\mu_r \lesssim 28 \text{ mag arcsec}^{-2}$ (Hood et al. 2018).

2.1. LERG Selection

In addition to the optical data, radio observations are required in order to detect and classify LERGs. To this end we select our LERGs from the Best & Heckman (2012) catalogue of 18,286 radio galaxies in SDSS. This catalogue is the result of cross matching observations from SDSS, the National Radio Astronomy Observatory (NRAO) Very Large Array (VLA) Sky Survey (NVSS, Condon et al. 1998), and the VLA Faint Images of the Radio Sky at Twenty cm survey (FIRST, Becker et al. 1995; Best et al. 2005b; Best & Heckman 2012). The Best & Heckman (2012) catalogue segregates radio galaxies into those where the radio emission is the result of star formation from those where it is the result of an AGN. This is based on a combination of stellar mass, 1.4 GHz luminosity, $H\alpha$ luminosity, 4,000 Å break strength, and emission line diagnostics (Baldwin et al. 1981; Kauffmann et al. 2003a; Best

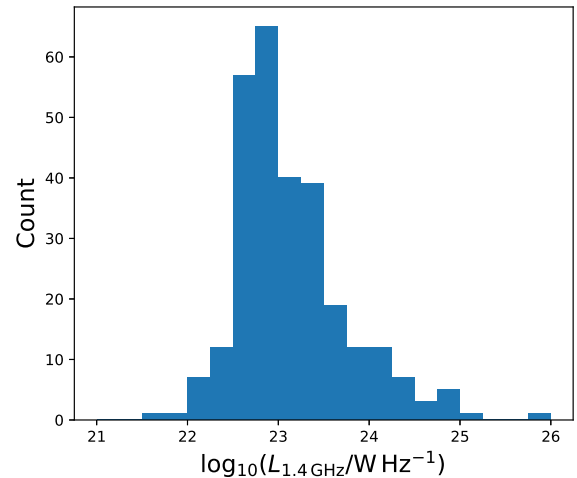


Figure 1. The distribution of 1.4 GHz luminosities of our LERG sample.

et al. 2005b; Kauffmann et al. 2008; Best & Heckman 2012). Radio-loud AGN are then further classified into high- or low-excitation sources based on the host galaxy spectrum. For high-quality spectra with many observed emission lines, the excitation index of Buttiglione et al. (2010) is used for the purpose. For poorer quality spectra, or those with intrinsically fewer emission lines, a more simple $\text{EW}_{[\text{O III}] \lambda 5007} > 5 \text{ \AA}$ criteria is used to segregate HERGs and LERGs. For a full description of the radio galaxy classification used in the construction of their catalogue, the reader is directed to Best & Heckman (2012), and Best et al. (2005b).

The nature of LSB astronomy necessitates that only local-Universe galaxies can be included in our analysis. For this reason we select LERGs from the Best & Heckman (2012) catalogue with $z < 0.07$ (Kaviraj 2014b). The LERGs selected cover four orders of magnitude in radio luminosity, $21.7 < (L_{1.4\text{GHz}}/\text{WHz}^{-1}) < 25.8$, with a median luminosity of 10^{23} WHz^{-1} at 1.4 GHz. The distribution of our LERG radio luminosities is shown in Figure 1. Additionally, their distribution in redshift, stellar mass, group/cluster halo mass (where applicable), and colour are shown alongside the control sample (described below) distributions in Figure 3.

2.2. The Main Control Sample

2.2.1. Redshift and stellar mass

In order to assess whether or not there is an excess of tidal remnants around LERGs, a suitable control sample of radio-quiet galaxies must be constructed. To this end we match each LERG to galaxies without a radio-detection on the basis of redshift and stellar mass. Stellar mass estimates for both the LERGs and control

galaxy candidates are taken from the MPA/JHU value added catalogue for SDSS DR7¹ (Tremonti et al. 2004). These stellar masses are calculated adopting the method of Kauffmann et al. (2003b) but using the SDSS *ugriz* photometry of the source rather than spectral indices². We require that control galaxy candidates have:

- $z_{\text{control}} = z_{\text{LERG}} \pm < 0.01$
- $M_{*,\text{control}} = M_{*,\text{LERG}} \pm < \sigma_{M_*}$,

Where σ_{M_*} is obtained from the 16th and 84th percentiles of the probability distribution of the stellar mass estimates.

Photometrically derived stellar mass estimates can be biased in AGN due to the excess short-wavelength radiation from the accretion disk (this is especially true in unobscured AGN, Gordon et al. 2017), and indeed powerful radio galaxies often exhibit an excess of ultraviolet light (Tadhunter et al. 2002). However, given that LERGs are inefficiently accreting AGN, and thus do not have an optically thick, luminous accretion disk, contamination of the broadband optical photometry from the AGN should be insignificant. Furthermore, Kauffmann et al. (2008) demonstrated that for the Best et al. (2005b) sample of radio galaxies in SDSS, which covers a radio power range inclusive of our LERG sample’s $L_{1.4\text{GHz}}$ distribution, such a UV excess was consistent with originating from a young stellar population rather than from AGN contamination. Thus, stellar mass estimates for our LERG population should be as reliable as the stellar mass estimates for galaxies not hosting an AGN. Moreover, obtaining stellar masses from the MPA/JHU catalogue for our LERGs is consistent with previous works (Smolčić 2009; Best & Heckman 2012).

2.2.2. Large-scale structure

Beyond controlling for just stellar mass and redshift, we require that each LERG is matched with control galaxies in the same type of large-scale environment. To determine this we use the Yang et al. (2007) SDSS group catalogue for DR7. For those galaxies within this catalogue we use the halo mass, M_{180} , of the structure the galaxy is located in to characterise its environment, considering halos of $M_{180} > 10^{12.5} M_{\odot}$ to be a group, and halos more massive than $10^{14} M_{\odot}$ to be a cluster (Gordon et al. 2018b; Barsanti et al. 2018; Lofthouse et al.

2018). Yang et al. (2007) observe that the error in obtained halo masses for their catalogue vary between 0.2 and 0.35 dex, and remain above 0.25 dex for the mass range $12.2 \lesssim \log_{10}(M_{180}/M_{\odot}) \lesssim 14.6$ (see their Fig. 7). By matching halos with $\Delta M_{180} < 0.25$ dex, we are thus matching them to other halos of reliable comparable mass. Galaxies in halos of $M_{180} < 10^{12.5} M_{\odot}$ are treated as field galaxies and are matched only with other field galaxies.

Given the known tendency for LERGs to be found within dense environments (e.g. Ching et al. 2017), it may be possible that our LERGs observed to lie within large scale structures may be more accurately assigned to groups and clusters than their controls. That is to say, the probability of group or cluster membership for a galaxy found by the group finding method of Yang et al. (2007) may be higher if the galaxy is a LERG. To determine if such an effect exists within our selected LERG and control populations, we calculate the *C*-statistic of Smith et al. (2004) for each of our selected galaxies. The *C*-statistic is a measure of likelihood of a galaxy to be associated with a particular structure and is given by

$$C = \frac{(cz_{\text{galaxy}} - cz_{\text{cluster}})^2}{\sigma^2} - 4 \log_{10} \left(1 - \frac{R}{R_{\text{cluster}}} \right), \quad (1)$$

where, z_{galaxy} is the galaxy redshift, z_{cluster} is the median cluster (or group) redshift, σ is the cluster velocity dispersion, R is the projected separation of the galaxy from the cluster centre, and R_{cluster} is the cluster radius. In Figure 2 the distributions of this statistic for both our LERG and control samples associated with large scale structures are shown to be similar for both populations. Additionally, these distributions fail to be separated by Kolmogorov-Smirnov (KS) testing, providing a *p*-value of 0.4, demonstrating that there is no significant difference in the accuracy of group and cluster membership assignment between our LERG and control populations.

For each LERG we select up to six (where possible) control galaxies satisfying these criteria, prioritised by closest match in stellar mass. Where fewer than six control galaxies for a LERG are found, as many controls as possible that satisfy the matching criteria are selected³. This results in a parent sample of 1,648 control galaxies matched to 284 LERGs. Whilst we do not control based on galaxy colour, we note that the *g* – *i* colour distribution of our control sample is similar to that of our LERGs. This, alongside the LERG and control sample

¹ <https://wwwmpa.mpa-garching.mpg.de/SDSS/DR7/>

² A comparison to the SDSS stellar masses obtained by Kauffmann et al. (2003b), who used spectral features rather than broad band photometry, can be found at https://wwwmpa.mpa-garching.mpg.de/SDSS/DR7/mass_comp.html

³ The weighting of controls in such cases is discussed in-depth in Section 5.1.3.

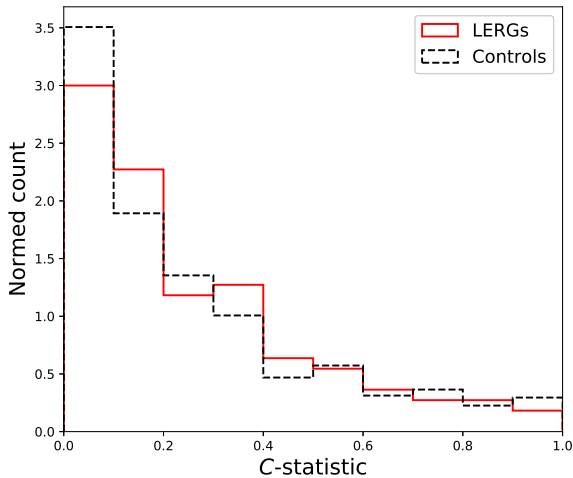


Figure 2. The distribution of calculated C -statistic (Smith et al. 2004) values for our selected galaxies assigned membership to large-scale structures, where a lower value is associated with increased likelihood of membership. The red solid line shows the C distribution for the LERG population, and the black dashed line for the control population.

distributions of redshift, stellar mass and, for non-field galaxies, halo mass are shown in Figure 3.

2.3. Controlling for Galaxy Morphology

Ideally the control sample should also be matched to the LERG population on morphology. To this end, we obtain morphological information for our selected LERGs and controls as part of the ongoing Galaxy Zoo project (Lintott et al. 2008; Willett et al. 2013; Walmsley et al. in prep). The current iteration of Galaxy Zoo⁴ has the advantage of using colour images from DECaLS rather than the SDSS images used by previous versions of the project. Whilst Galaxy Zoo provides data on various subtle morphological parameters (e.g., number of spiral arms, Willett et al. 2013), for this work we wish only to know if a galaxy is early- or late-type.

To classify our galaxies as early- or late-type, we require at least 95 % confidence in voting for one particular morphology over another, where the confidence limits assume binomial errors (Cameron 2011) and are calculated from the raw number of votes for each answer. The initial question in the current Galaxy Zoo workflow (to be described in full in Walmsley et al. in prep, and based on Willett et al. 2013), is concerning the broad morphology of the galaxy, i.e. is it *smooth and rounded* (early-type) or *disk or featured* (late-type).

To classify our galaxies as either early- or late-type we require that the 95 % confidence intervals for these two answers do not overlap. Additionally we require that, for the favoured classification, the voting indicates a majority verdict (i.e. > 50 % of the total votes) at greater than 80 % confidence.

Using this method, we obtain reliable morphologies for 216 of our LERGs and 1,138 of the control galaxies. Of our LERGs with reliable classifications 91 % are classified as early-type, compared to 67 % of the control sample. Ensuring the control galaxy morphology is the same as the LERG morphology (in addition to satisfying the other control parameters described in Section 2.2) results in a morphologically-matched subset of 191 LERGs and 657 controls. Whilst our primary analysis is conducted on the overall LERG and control samples, this morphologically-matched subset provides a benchmark to compare our results with and thus ensure that our findings are not biased by galaxy morphology.

3. ANALYSIS OF DECALS IMAGING

For each of the LERGs and control galaxies, a DECaLS DR5 r -band cutout, measuring 200×200 kpc at the redshift of the target galaxy, is obtained. There are some small regions in DECaLS DR5 where the imaging in all three bands is incomplete, resulting in a small fraction (~ 1 %) of our data selection not having DECaLS DR5 r -band imaging. The images were processed using an arcsinh stretch, allowing the fainter details to be seen whilst preserving the brighter features better than is possible using a logarithmic scaling (Lupton et al. 2004). The resultant ‘postage stamp’ image is then uploaded to the Zooniverse project builder⁵, where the images are presented in a random order for blind classification. We show examples of these in Figure 4 alongside SDSS colour images for comparison. Each image presented had four possible options to vote for, of which only one could be selected:

- ‘*No disturbance*’ required that there be no obvious asymmetries to the low- or high-surface brightness features, no apparent tidal tails or shocks (e.g., column one of Figure 4).
- ‘*Minor disturbance*’ was selected should an image have features affecting the LSB morphology, e.g., halo shells or faint tidal streams (see column two of Figure 4).

⁴ <https://www.zooniverse.org/projects/zookeeper/galaxy-zoo>

⁵ <https://www.zooniverse.org/lab>

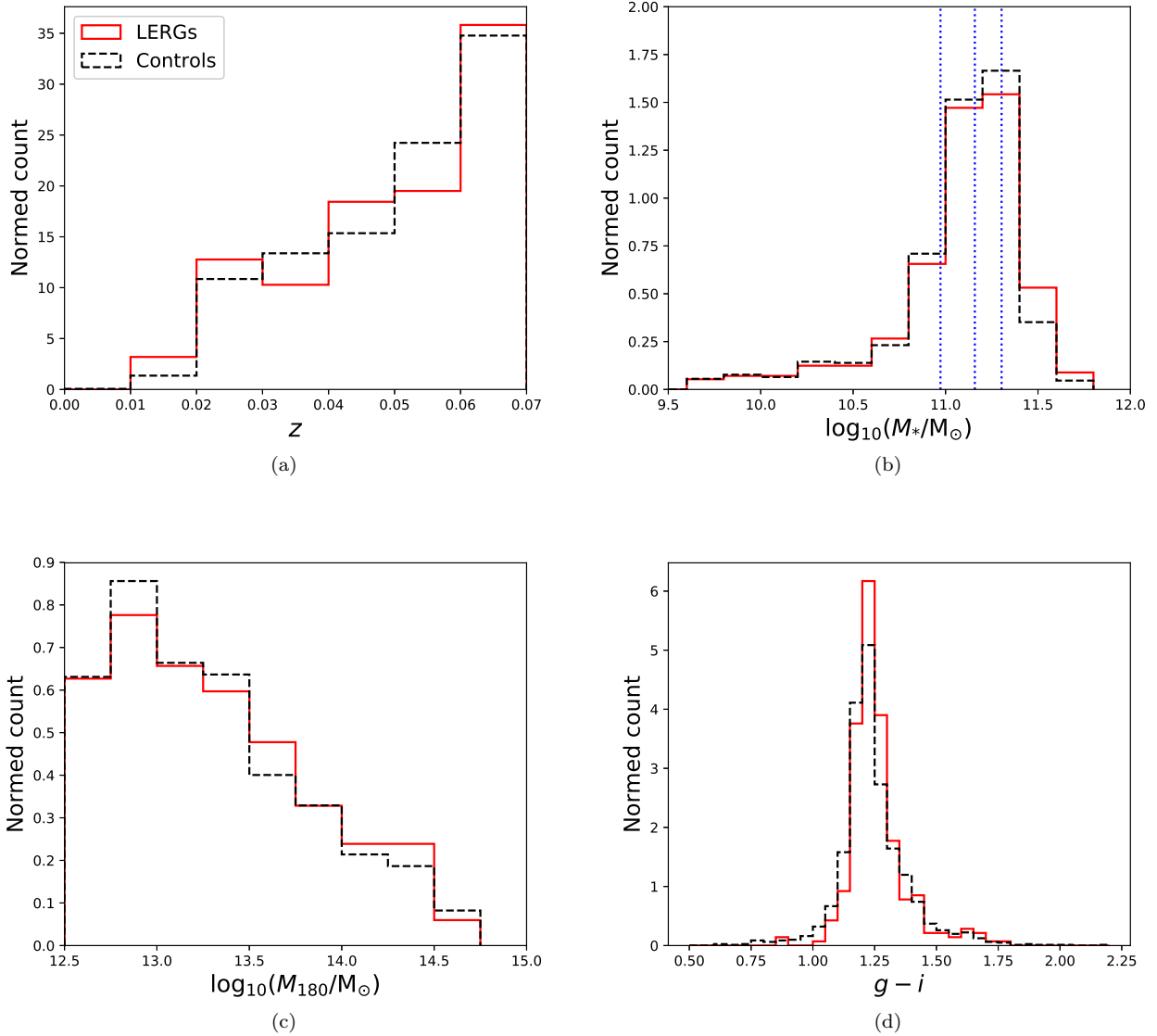


Figure 3. The similar distributions of the redshift (panel a), stellar mass (panel b), group halo masses (panel c), and the $g - i$ colours (panel d) of the LERG and control samples. The red solid line represents LERGs and the black dashed line represents the control population. The blue dotted lines shown in panel b highlight the 25th, 50th, and 75th percentiles of the LERG stellar mass distribution, used for mass subsetting in our analysis.

- ‘*Major disturbance*’ was dependent on there being clear disruption to the high surface brightness morphology of the galaxy, frequently with second similar size galaxy involved (as shown in column three of Figure 4).
- ‘*Bad data*’ was included as an option should the image quality prevent classification (e.g. due to bad stitching or artefacts).

The classification of the images is a subjective task for which there may be variance between individual voters. However, as each person’s voting was conducted independently, individual analyses should be self consistent. In total seven volunteers from the coauthorship were involved in the classification process, providing 10,796 votes for a total of 1,904 images, with each image having been classified by between five and seven of these classifiers.

4. RESULTS

4.1. The Prevalence of Tidal Features in LERGs

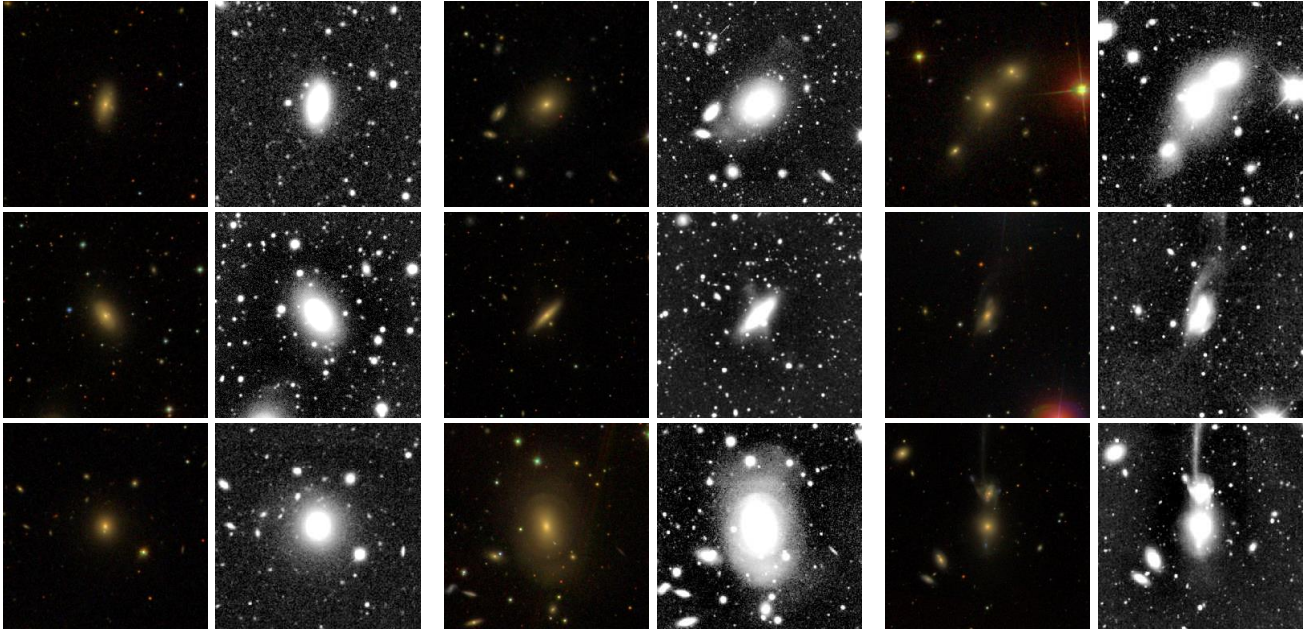


Figure 4. Example images from SDSS and DECaLS of galaxies of different classifications from our analysis, three examples of each classification are given - one per row. Split into 3 columns, each 2 panels wide, the left panel in a column shows an SDSS standard-depth colour image of the galaxy, whilst the right panel in a column shows the deeper DECaLS DR5 r -band image used in this work. The left-most column shows galaxies classified as having no morphological disturbance, the middle column shows minor morphological disturbances, and the right-most column shows major morphological disturbances. Note how for the minor, and even the major, disturbances the DECaLS images readily show LSB features that are barely noticeable or indeed totally absent in the SDSS imaging.

Images of 282 LERGs and 1,622 control galaxies were analysed. The votes from each classifier were stacked and the fraction of each of the possible classifications calculated. Uncertainties are estimated by bootstrapping the result 10,000 times, and using the 16th and 84th percentiles of the resultant distribution as the lower and upper errors⁶. We find that the overall merger fractions of LERGs and controls are broadly consistent, with $28.7 \pm 1.1\%$ of LERGs and $27.3 \pm 0.5\%$ of controls observed to have tidal features. In the LERG sample, major tidal features are found in $7.2 \pm 0.6\%$ of the population, while $21.4^{+1.1}_{-1.0}\%$ of LERGs are observed to have minor morphological disturbances. For the control sample, $5.0 \pm 0.2\%$ have major morphological disturbances, and $22.3 \pm 0.5\%$ have minor tidal remnants. Whilst minor disruption is seen in the LSB morphologies of LERGs and controls at similar rates, there is an excess of major tidal features observed in the LERG population at $> 3\sigma$ confidence. The fractions of LERGs and controls with major and minor tidal features are shown in Table 1.

⁶ Using P_{16} and P_{84} as the error estimates here, rather than the standard deviation, allows for potential variance in the symmetry of the bootstrapped distribution when calculating the 68% confidence interval.

Our morphologically-matched subsample shows similar results to the analysis conducted without controlling for morphology. For these LERGs we report that $5.3 \pm 0.7\%$ and $17.4 \pm 1.1\%$ are associated with major and minor tidal features respectively. Their control sample shows $3.4 \pm 0.3\%$ observed to have major tidal remnants, and $20.8 \pm 0.7\%$ with minor morphological disruption. This suggests, with $\sim 2.5\sigma$ confidence, that the tendency for LERGs to be more likely than controls to have substantial tidal features is not dependent on the galaxy morphology. We attribute the systematically lower fractions of galaxies with tidal features in this morphologically-matched subsample to the need for a reliable morphological classification, on which tidal remnants may naturally impact. Table 2 shows the fractions of LERGs and their morphologically matched controls with tidal features

4.2. Are Tidal Features More Common in Low- or High-Mass LERGs?

To investigate whether the fraction of LERGs having tidal features evolves with stellar mass, we repeat the analysis for the outer quartiles and the interquartile range of the stellar mass distribution of our LERG host galaxies. Doing this we find that the excess of LERGs having major tidal features is driven by ‘lower-

Table 1. Fraction of LERGs and control galaxies with tidal features for the full sample and various mass and environment subsets. The first column defines the subset of the population analysed. The second and third columns show the fractions of galaxies with minor tidal remnants, and the final two columns show the fractions of galaxies with major tidal disruption. The top row shows the entire population. The next three rows show stellar mass subsets of the population, with ‘*low mass*’ defined by the lowest quartile of the stellar mass distribution of our LERG sample, ‘*intermediate mass*’ by the interquartile range, and ‘*high mass*’ by the upper quartile (see panel b of Figure 3). Environmental subsets of the population are shown in the lowest three rows.

Galaxy subset	LERG - minor	Control - minor	LERG - major	Control - major
	[%]	[%]	[%]	[%]
<i>All</i>	$21.4^{+1.1}_{-1.0}$	$22.3^{+0.5}_{-0.5}$	$7.2^{+0.6}_{-0.6}$	$5.0^{+0.2}_{-0.2}$
<i>Low mass</i>	$15.9^{+1.8}_{-1.8}$	$14.3^{+0.8}_{-0.8}$	$10.0^{+1.5}_{-1.5}$	$3.2^{+0.4}_{-0.4}$
<i>Intermediate mass</i>	$21.6^{+1.5}_{-1.5}$	$24.0^{+0.7}_{-0.7}$	$6.2^{+0.9}_{-0.9}$	$4.6^{+0.3}_{-0.3}$
<i>High mass</i>	$26.4^{+2.0}_{-2.2}$	$27.3^{+1.0}_{-1.0}$	$6.5^{+1.2}_{-1.2}$	$7.8^{+0.6}_{-0.6}$
<i>Field galaxies</i>	$19.8^{+1.3}_{-1.5}$	$21.3^{+0.6}_{-0.6}$	$7.8^{+1.0}_{-1.0}$	$4.6^{+0.3}_{-0.3}$
<i>Group galaxies</i>	$25.5^{+1.7}_{-1.7}$	$23.5^{+0.7}_{-0.7}$	$6.7^{+0.9}_{-1.1}$	$5.4^{+0.4}_{-0.4}$
<i>Cluster galaxies</i>	$8.8^{+2.9}_{-2.9}$	$23.0^{+2.0}_{-2.0}$	$5.9^{+2.0}_{-2.0}$	$6.1^{+1.1}_{-1.1}$

Table 2. Fraction of LERGs and control galaxies with tidal features for the morphologically-matched sample and various mass and environmental subsets. The layout is as for Table 1.

Mass bin	LERG - minor	Control - minor	LERG - major	Control - major
	[%]	[%]	[%]	[%]
<i>All</i>	$17.4^{+1.1}_{-1.1}$	$20.8^{+0.7}_{-0.7}$	$5.3^{+0.7}_{-0.7}$	$3.4^{+0.3}_{-0.3}$
<i>Low mass</i>	$12.4^{+2.0}_{-2.0}$	$10.8^{+1.1}_{-1.1}$	$5.6^{+1.6}_{-1.6}$	$1.4^{+0.4}_{-0.4}$
<i>Intermediate mass</i>	$19.7^{+1.8}_{-1.6}$	$23.4^{+1.0}_{-1.0}$	$5.2^{+0.9}_{-0.9}$	$3.4^{+0.4}_{-0.4}$
<i>High mass</i>	$17.1^{+2.4}_{-2.4}$	$24.8^{+1.5}_{-1.4}$	$5.3^{+1.6}_{-1.6}$	$5.5^{+0.9}_{-0.7}$
<i>Field galaxies</i>	$14.5^{+1.6}_{-1.6}$	$20.3^{+1.0}_{-1.0}$	$7.7^{+1.2}_{-1.2}$	$2.6^{+0.4}_{-0.4}$
<i>Group galaxies</i>	$22.2^{+1.9}_{-1.9}$	$20.9^{+1.1}_{-1.0}$	$3.4^{+0.8}_{-0.8}$	$3.6^{+0.4}_{-0.4}$
<i>Cluster galaxies</i>	$5.7^{+2.9}_{-2.9}$	$23.9^{+3.3}_{-3.3}$	$1.4^{+1.4}_{-1.4}$	$8.9^{+2.2}_{-2.2}$

mass’, $\log_{10}(M_*/M_\odot) \lesssim 11$, galaxies. When only the lower quartile of our LERG stellar mass distribution, $\log_{10}(M_*/M_\odot) < 10.97$, is considered, $10.0 \pm 1.5\%$ of LERGs are observed to have these major remnants compared to $3.2 \pm 0.4\%$ of control galaxies. That is to say, an excess of major tidal features in LERGs at lower stellar masses is seen with $> 4\sigma$ confidence. When the morphologically-matched sample is considered this excess is seen with $> 2.5\sigma$ confidence for this mass range, and rising to $> 3.5\sigma$ confidence for $\log_{10}(M_*/M_\odot) < 11.16$ (i.e. if we stack the two lowest mass quartiles). For the interquartile range and the upper quartile of the stellar mass range, this excess of major mergers in the LERG population is not observed to a significant level. This evolution with stellar mass

for LERGs to be associated with major tidal features is shown in Figure 5, and in Table 1.

When compared to the morphologically-matched control subset, LERGs are observed to have a deficit of minor tidal remnants with $> 2.5\sigma$ confidence in the highest mass bin, i.e., $\log_{10}(M_*/M_\odot) > 11.3$. At lower masses, or without controlling for morphology, the fraction of LERGs with minor tidal features is consistent with the control sample. As with the observations of major tidal features, the fractions of LERGs and control galaxies in different mass bins with minor tidal features is shown in Figure 5 as well as in Table 2.

4.3. The Influence of Galaxy Environment on the Likelihood of LERGs to Have Tidal Features

We additionally compared the fraction of LERGs and control galaxies with tidal features within different large-

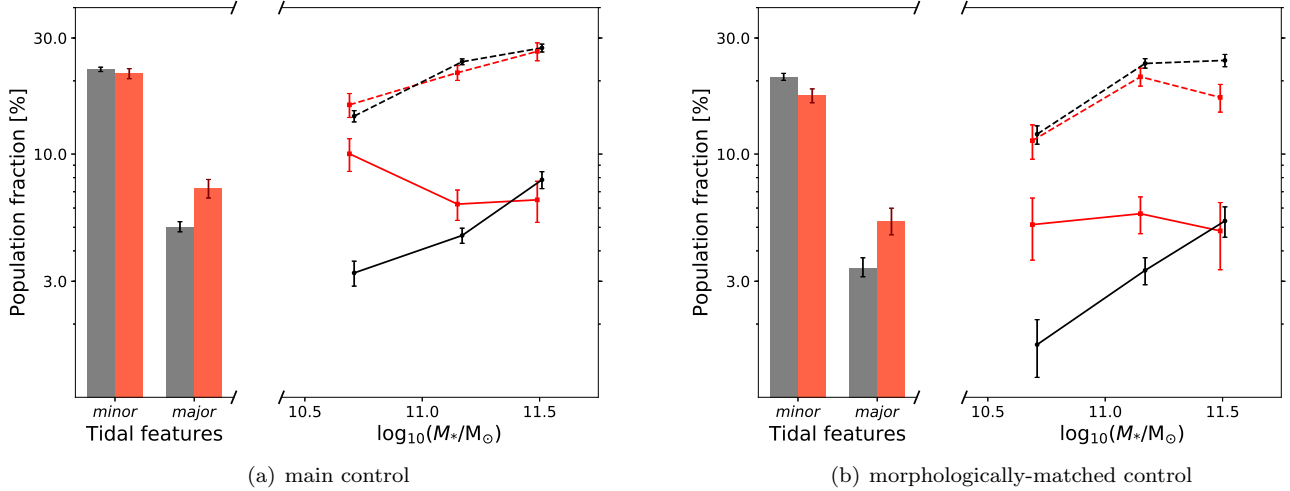


Figure 5. The fractions of LERGs and controls showing minor and major tidal features, with panel a used to show the larger main control sample, and panel b showing the morphologically-matched subset. The overall fractions are shown in the bar plots on the left, while the fractions in different stellar mass bins, specifically the interquartile range and the two extreme quartiles, are shown on the right. Here the dashed lines show the minor tidal feature fractions, and the solid lines show the fractions of galaxies with major tidal features. The LERG and control populations on the right are horizontally offset from each other for clarity. The excess of LERGs undergoing major interactions is clearly shown to be driven by the low mass end of our sample. Red is used to represent the LERG population and grey/black represents the control sample.

scale environments, i.e, whether the galaxy is in the field, a group, or within a cluster. For the purposes of differentiating between groups and clusters, we segregate these structures at halo masses of $10^{14} M_{\odot}$. That is to say structures with $M_{180} < 10^{14} M_{\odot}$ are classed as groups, and those with $M_{180} > 10^{14} M_{\odot}$ are considered to be clusters (Lofthouse et al. 2018; Barsanti et al. 2018). Galaxies in very low mass groups, $M_{180} < 10^{12.5} M_{\odot}$, are treated as field galaxies (Gordon et al. 2018b). The merger fractions of LERGs and controls in these different environments are presented in Tables 1 and 2, and in Figure 6, as well as being described below.

4.3.1. *Field galaxies*

When only field galaxies are considered the excess of major tidal features in LERGs is detected at the $\sim 3\sigma$ level, with $7.8 \pm 1.0\%$ of LERGs having such features compared to $4.6 \pm 0.3\%$ of the control sample. As with the whole population this excess is greatest in the lower mass regime of our sample. To maintain a sizeable population whilst subsetting on multiple parameters, here we simply look at galaxies above and below the median stellar mass of our sample, $\log_{10}(M_*/M_{\odot}) = 11.16$. We note that the LERG excess of major tidal features is seen at $> 2.5\sigma$ in the lower mass field population, compared to $< 2\sigma$ for the higher mass galaxies. Requiring the control population to be matched on morphology does not influence this observation. Morphology-matching does show that whilst $20.3 \pm 1.0\%$ of control galaxies have

minor tidal features, only $14.5 \pm 1.6\%$ of LERGs exhibit such morphological disturbances, a $\sim 3\sigma$ deficit. This appears to be driven by stellar mass with this excess becoming insignificant at $\log_{10}(M_*/M_{\odot}) < 11.16$.

4.3.2. *Galaxy groups*

In galaxy groups when the whole mass distribution is considered, no significant trends are observed for the fraction of LERGs with tidal features (either major or minor) to differ from control population. However, as with the overall analysis, and the analysis on field galaxies, differences in the fractions of control galaxies and LERGs with major morphological disruptions are noticed in different mass regimes. In the lower quartile of the stellar mass distribution the excess of major morphological disturbances in LERGs is seen at the $\sim 3\sigma$ level, with $14.0 \pm 0.4\%$ of LERGs having such disruption to their morphology compared to $3.2 \pm 0.9\%$ of control galaxies. At the highest stellar masses this trend is inverted. As few as $2.6 \pm 1.0\%$ of LERGs with $M_* > 10^{11.3} M_{\odot}$ in groups have major morphological disruption in comparison to $7.6 \pm 0.9\%$ of the control sample, a $> 3.5\sigma$ deficit. The requirement for the control sample to be matched to the LERGs on morphology has no impact in these results.

4.3.3. *Galaxy clusters*

In galaxy clusters the fraction of LERGs with major tidal features is consistent with the control popu-

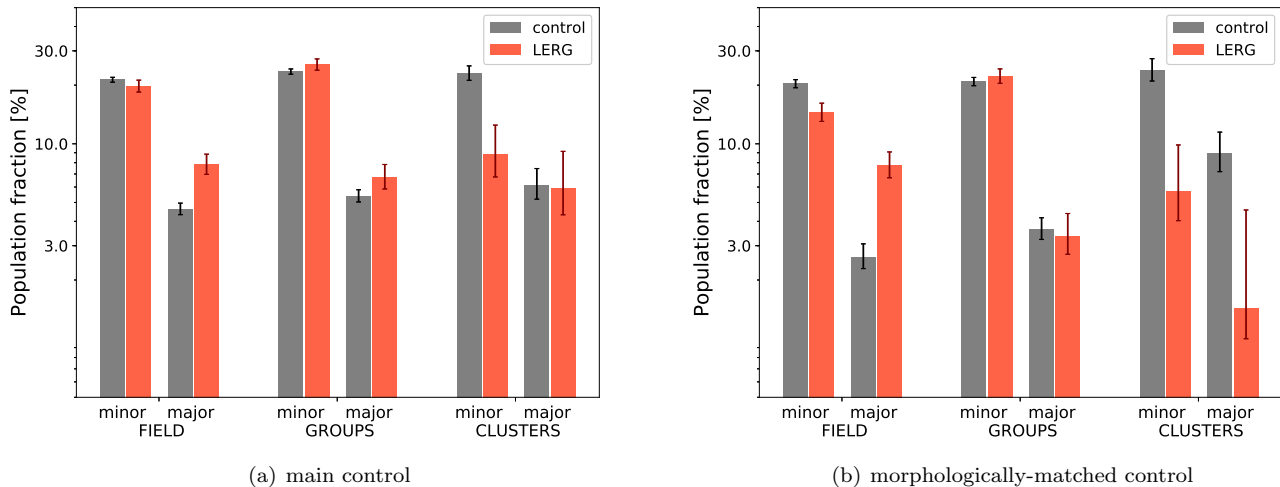


Figure 6. The effect of galaxy environment on the LERG merger fractions with respect to the control population. For each environment (field, groups or clusters) the minor merger fractions for the LERGs (red) and controls (grey) are shown on the left, whilst the major merger fractions are shown on the right. The left-hand panel shows the full LERG sample and control, whilst the right-hand panel shows the morphologically-matched subset.

lation. Even at in the lower mass galaxies, where an excess of LERGs with major tidal features appears to be strongest, such an effect is only detected with $< 2\sigma$ confidence. Furthermore a deficit in LERGs with major tidal features is noted at $> 2.5\sigma$ confidence when LERG morphology is controlled for.

Considering the galaxies with minor disturbances to their morphology, $8.8 \pm 2.9\%$ of LERGs in clusters display such features. This presents a $\sim 4\sigma$ deficit relative to the control population, where $23.0 \pm 2.0\%$ of galaxies show minor morphological disruption. No significant variance of this result is observed with changes in the stellar mass range analysed, or with the requirement for LERG morphology to be matched in the control.

5. INTERPRETATION OF OBSERVATIONS AND DISCUSSION

5.1. Potential Sources of Bias

5.1.1. Spectroscopic targeting and completeness

Both our LERG and control populations are selected from DR7 of the SDSS spectroscopic catalogue (Abazajian et al. 2009). One of the limitations of this catalogue is the effect of fibre-collisions that prevent any two spectroscopic fibres on the same plate being positioned closer together than $55''$ (Strauss et al. 2002; Patton & Atfield 2008). Consequently, spectroscopic completeness is impeded in regions of the sky with a high target density (e.g., Yoon et al. 2008; Robotham et al. 2010; Gordon et al. 2017, 2018a).

Due to fibre-collision induced incompleteness, SDSS targeting algorithms prioritise some targets ahead of

others. Specifically with respect to this work, objects identified as quasar candidates are prioritised ahead of the main galaxy sample (Strauss et al. 2002; Blanton et al. 2003). One might expect a low- z sample of LERGs, such as ours, which are predominantly red, extended sources, not to be selected as quasar candidates based on their optical photometry. However, quasar candidates in SDSS are identified both by their optical photometry, and the presence of radio emission via cross-matching with FIRST (Richards et al. 2002). Should our LERG sample have a higher targeting priority in dense regions of sky, then this could potentially bias the comparative merger fractions we observe in the LERGs and control galaxies.

The entirety of our LERG sample and more than 99% of our control population are included in the SDSS DR7 main galaxy sample. Consequently, reanalysing our observations limited to just SDSS main sample galaxies has no impact on our results. As an additional check to ensure there is no difference in the density of potential spectroscopic targets close proximity to our LERG and control samples, we analyse the number of neighbours observed in the SDSS DR7 photometric catalogue, and with $14.5 < r < 17.77$, around both populations. We find that the number of neighbours within $55''$ is consistent for the LERG and control populations, with an average of 1.18 ± 0.06 and 1.10 ± 0.03 neighbours respectively. Reducing the search radius to $25''$, i.e. less than half the fibre-collision limit, we observe an average of 0.33 ± 0.06 neighbours per LERG, and 0.29 ± 0.03 neighbours per control galaxy. We are thus confident

that our merger fractions are not biased due to differences in on-sky target densities.

5.1.2. Treating low mass halos as the field

In matching the LERG and control samples on large-scale environment, we chose to consider galaxies within halos with $M_{180} < 10^{12.5} M_{\odot}$ as being in the field rather than in groups. However, given that LERGs are frequently observed in over-dense environments, it may be the case that LERGs classified in the field may be more likely than their control galaxies to be a member of a low mass group. If this were the case, it could explain the increased fraction of our ‘field’ LERGs that exhibit tidal features relative to the control sample.

In our LERG and control populations classified as being field galaxies, $18.9^{+3.6}_{-2.8} \%$ and $19.9^{+1.4}_{-1.3} \%$ respectively reside in these low mass halos. Additionally, for the galaxies in the low mass groups that we have classified as the field, the tendency for LERGs to be found in over-dense environments may translate to them being in higher mass halos than the control galaxies. Comparing the halo mass distributions of these low mass groups shows no difference between the groups hosting LERGs and control galaxies, with a KS derived p -value of 0.98 (see also Figure 7). Furthermore, we re-conduct our analysis of the field galaxy population using only those galaxies hosted by low mass groups. Here we find no excess of LERGs showing major morphological disturbances, with $2.8 \pm 1.4 \%$ of LERGs displaying such features compared to $3.5^{+0.9}_{-1.1} \%$ of their controls. This demonstrates that the observation of excess major tidal disruption in field LERGs is not driven by LERGs residing in the lowest mass groups.

5.1.3. Weighting of the control sample

In conducting this work each control galaxy was given an equal weighting, and, in order to maximise the sample size, all possible control galaxies were used for any particular analysis. As stated in Section 2, where possible we have selected six control galaxies per LERG. However, in $\sim 10 \%$ of cases fewer than six control galaxies could be found for a LERG, and in these cases as many controls as can be found that satisfy the matching criteria are included. Thus, there is the potential for this variance in the number of control galaxies available per LERG to influence our results. Given that just $\sim 10 \%$ of LERGs have fewer than six control galaxies selected, then a quick test of such an effect would be to restrict our analysis to those LERGs with the maximum number of available control galaxies (252 LERGs and 1512 control galaxies). Doing this, we find no substantial changes to our results, with all of our statistically significant observations remaining above 3σ confidence.

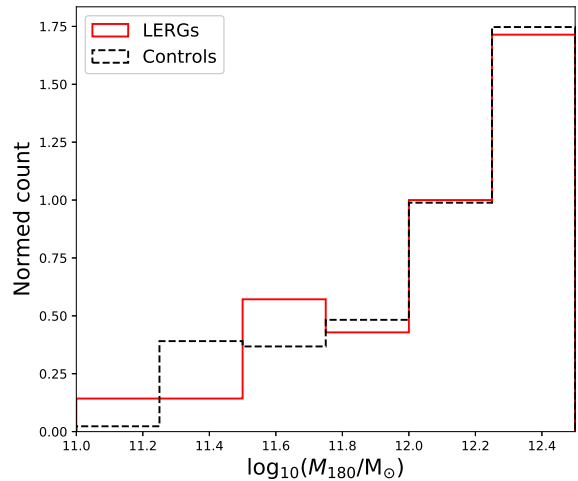


Figure 7. The distributions of halo mass estimates for galaxies assigned to groups with $M_{180} < 10^{12.5} M_{\odot}$, and thus treated as field galaxies in our analysis. The red solid line shows the low mass groups containing LERGs, and the black dashed line the groups containing their control galaxies. Applying a two-sample KS test to these distributions returns $p = 0.98$, consistent with both distributions being drawn from the same parent population. Halo mass estimates obtained from the Yang et al. (2007) SDSS group catalogue.

Our observations are thus not significantly affected by the limited number of control galaxies available to some of the selected LERGs.

5.1.4. Interpreting Merger Scale From Tidal Feature Intensity

Although we have classified galaxy images based on the intensity of their tidal features (should any be present), does this necessarily translate to merger intensity? In particular, in the very latest stages of major merger, even the most substantial of tidal features will dissipate over time as the two galaxies coalesce. Consequently a very late stage major merger may show only minor disruptions to the extended LSB morphology (for example see Figure 1. in Lotz et al. 2008).

The degeneracy in the origin of minor LSB features naturally complicates attributing them to either a minor or major merger. It is anticipated that straight tidal streams (e.g. as shown in column 2, row 2 of Figure 4) are the result of recent minor mergers rather than being older remnants (Duc et al. 2015). Furthermore, the shell structures seen around many early type galaxies (e.g. column 2, row 3 of Figure 4), and which we have classified as a minor morphological disruption, are frequently associated with low-intermediate mass ratio mergers (i.e with mass ratios in the range 1:3 to 1:10,

Kaviraj 2010; Duc et al. 2015; Pop et al. 2018; Kado-Fong et al. 2018). Considering these points in combination with the relatively short period during which a major merger exhibits only minor tidal features, any contamination of the minor disturbance sample by late stage major mergers should be minimal. For these reasons we adopt the approach that minor tidal features are the result of minor mergers, whereas major features are due to major mergers, typical of previous works involving classification of morphological disruptions in large samples of galaxies (e.g. Kaviraj 2014a,b; Morales et al. 2018).

5.2. The Influence of Major Mergers on LERG Evolution

5.2.1. LERGs can evolve from a major merger

Our observations show a $> 3\sigma$ excess of LERGs undergoing major mergers in comparison to a control sample, and this excess is shown to be strongest ($> 4\sigma$ confidence) at lower stellar masses. This mass trend of the LERG major merger fraction persists in all environments. Although only seen at less than 2σ confidence for galaxies within clusters, even here it is at the lower end of the stellar mass distribution of our sample where any potential excess of major mergers in LERGs is noted. This demonstrates that major mergers can result in the production of a LERG, particularly in field and group galaxies, and at stellar masses below $\sim 10^{11} M_{\odot}$.

Prior observations have shown mixed results with regard to the likelihood of major mergers being involved in the evolution of a galaxy into a LERG. For galaxies with $10 < \log_{10}(M_*/M_{\odot}) < 12$, Sabater et al. (2013) demonstrated that LERGs were more likely than non-LERGs to be involved in one-on-one interactions with other galaxies. On the other hand, Ellison et al. (2015) demonstrated that LERGs were no more likely to be in a close galaxy pair than other similar galaxies after controlling for galaxy properties and large scale structures, indicating that major mergers are not the primary trigger mechanism for LERGs. In post-merger galaxies showing tidal remnants however, Ellison et al. (2015) do observe a slight, but insignificant, excess of LERGs relative to their control sample, potentially consistent with our results and those of Sabater et al. (2013). Furthermore, while Ellison et al. (2015) do not state the stellar mass range of their sample, their Figure 1 demonstrates that their LERGs in pairs are drawn from a population dominated by galaxies with $M_* > 10^{11} M_{\odot}$. It is thus possible that their work either does not include the low-mass galaxies that are driving the excess of major mergers in LERGs that we observe, or that these constitute a less substantial fraction of their LERG sample.

Whilst our observations show that major mergers clearly play a role in the triggering of some low-mass LERGs, even in this mass regime only $\sim 10\%$ of LERGs are currently experiencing such large-scale interactions. Major mergers therefore are not the dominant pathway to LERG activity. Hence, for the remaining $\sim 90\%$ of the low-mass LERG population, and for LERGs of higher stellar masses, other trigger mechanisms must be invoked.

5.2.2. Are major merger-driven low-mass LERGs the progenitors of HERGs?

The relatively low accretion rates associated with the nuclear activity in LERGs suggests that while major mergers can trigger low-mass LERGs, they don't directly fuel the black hole accretion within these galaxies. Rather the fuel source may be internal in origin, and that the accretion on to the supermassive black hole is induced by disk instabilities that are the result of pre- and in-merger gravitational perturbations (Bournaud et al. 2011; Nealon et al. 2015; Gatti et al. 2016; King & Nixon 2018). Such a mechanism has the potential to provide a more limited fuel supply than one might expect if the fuel originated from the donor component of the merging system, and hence explain the radiatively inefficient accretion mode observed. Moreover, radiatively inefficient black hole accretion may be just the first step of nuclear activity in these galaxies. It is reasonable to expect that even for efficiently accreting AGN there exists a short phase where the accretion rate is radiatively inefficient (Sabater et al. 2013). In this scenario, and if the merger is gas-rich, these merger-induced LERGs could then evolve into HERGs once the gas from the merger has fallen into the central engine providing the opportunity for more radiatively efficient accretion to take place. Such a process would be consistent with observations showing that Seyfert-like emission line ratios peak towards the end of the merger process (Carpinetti et al. 2012), and the association between mergers and powerful radio galaxies seen at high- z (Chibera et al. 2015).

These major merger-driven LERGs only represent a small fraction of the LERG population, and thus may not be typical of inefficiently accreting RLGN. The excess of major mergers in LERGs is seen most strongly at $\log_{10}(M_*/M_{\odot}) < 10.97$. It is interesting to note that while the median of our LERG stellar mass distribution is ~ 0.2 dex higher than this, the median stellar mass of HERGs in the Best & Heckman (2012) catalogue with $z < 0.07$ is $\log_{10}(M_*/M_{\odot}) \sim 10.9$, a difference of < 0.1 dex. In other words, the masses of LERGs where we observe the strongest excess of major mergers, are more typical of the HERG population than of the

broader LERG population, and, if merger-driven LERGs are potential progenitors of HERGs, one would expect their masses to be broadly consistent.

Should this description be an accurate representation of the physics at play, then low-mass LERGs associated with major merger systems represent young AGN, ostensibly at the point of trigger. These objects may thus be associated with relatively compact radio morphologies compared to LERGs that are not merger-driven, due to the jet having limited propagation time. Whilst we have made no attempt to do so in this work, a comparative analysis of the radio properties of merging versus non-merging LERGs presents a compelling opportunity for follow up. The next generation suite of high-resolution radio surveys such as the Very Large Array Sky Survey⁷ (VLASS, Lacy et al. in prep) should be well suited to such an analysis.

Such a smooth transition from radiatively inefficient to efficient accretion modes as the AGN evolves may of course be an oversimplification. It has been demonstrated that AGN can ‘flicker’ on and off over timescales $\lesssim 10^5$ years (Schawinski et al. 2015; King & Nixon 2015; Comerford et al. 2017). In radio quiet AGN, a drop in accretion rate below $\sim 0.01 \dot{M}_{\text{Edd}}$ would appear as an AGN being switched off. In RLAGN however, the presence of the relativistic jet ensures the galaxy is still detected as an AGN even at low accretion rates. This presents an alternative possibility to a steady LERG-to-HERG evolution, in that some LERGs may simply be the low-accretion phase of RLAGN flickering. However, in general, there are differences in both the stellar populations and radio properties of LERGs and HERGs (e.g. Baum et al. 1992; Buttiglione et al. 2010; Best & Heckman 2012). Consequently a comparison of these properties in major merger-driven LERGs and HERGs of similar mass will be required to test such a hypothesis, and is beyond the scope of this work.

5.3. The Role of Minor Mergers in the Evolution of LERGs

5.3.1. LERGs are not primarily fuelled by minor mergers

Given the requirement for a low Eddington scaled accretion rate in LERGs, minor mergers have been proposed as a potential pathway with which to introduce a limited fuel supply to the central engine (Kaviraj 2014a; Pace & Salim 2014; Ellison et al. 2015). Should this be the case, then LERGs are expected to have a higher fraction of tidal features associated with minor mergers than a control population (Kaviraj 2014a,b), and test-

ing this hypothesis is one of the principle aims of this work. We observe no such excess of minor mergers in our LERG sample, suggesting that these events do not play a substantial role in the fuelling of LERGs.

The validity of our test is dependent on a couple of assumptions regarding the visibility of the remnants of minor mergers. The merging system must be able to leave a detectable remnant with a surface brightness of $\mu_r < 28 \text{ mag arcsec}^{-2}$. Ji et al. (2014) investigated the visibility of tidal remnants resulting from different scale mergers involving a simulated galaxy with $\log_{10}(M_*/M_\odot) \sim 10.4^8$. Their work demonstrates that such a galaxy experiencing a $>1:10$ merger can produce tidal features with $\mu_r < 28 \text{ mag arcsec}^{-1}$. As the majority of satellites to massive galaxies are dwarfs with $M_* < 10^9 M_\odot$ (e.g., Loveday 1997; De Rijcke et al. 2006), then clearly a substantial number of very-minor mergers will go undetected. This may be particularly true for RLAGN given their increased number of satellites relative to inactive galaxies (Pace & Salim 2014). Furthermore, if the minor merger is to directly fuel the LERG, then the merger remnant must remain visible long enough for the gas from the donor galaxy to fall in to the central engine of the recipient. Based on the stellar population ages of AGN, this process is estimated to take several hundred Myr (Tadhunter et al. 2005; Bessiere et al. 2014; Shabala et al. 2017). At surface brightnesses of $\mu_r < 28 \text{ mag arcsec}^{-2}$, Ji et al. (2014) show that the remnant from a 1:6 mass ratio merger should be observable for $> 2 \text{ Gyr}$ post-merger. The visibility timescale of the merger remnant should therefore not prohibit the association of low to moderate mass ratio mergers with nuclear activity. Consequently we can make the statement that moderate to minor mergers are not the primary fuel supply for LERGs.

5.3.2. Minor mergers inhibit LERG activity in clusters

Although minor mergers do not preferentially trigger LERGs, it may not be accurate to say these events play no role in the evolution of an inactive galaxy into a LERG. We observe a significant (4σ) deficit of minor mergers in LERGs residing within clusters. Beyond just failing to contribute to the triggering of a LERG, these observations indicate that minor mergers may actually prevent galaxies within clusters from evolving in to LERGs.

Within clusters it has been shown that different regions of the cluster environment provide different opportunities for AGN fuelling (e.g. Haines et al. 2012; Pimb-

⁷ <https://science.nrao.edu/science/surveys/vlass>

⁸ This value is determined from the values of bulge stellar mass and disk stellar mass provided in Table 1 of Ji et al. (2014).

blet et al. 2013; Gordon et al. 2018b). The cluster core for instance may allow for cooling flows onto a galaxy, a widely hypothesised mechanism for LERG fuelling (e.g. Gaspari et al. 2013, 2017; Tremblay et al. 2016). The outer regions of clusters provide more opportunities for low-speed interactions such as mergers, and infalling galaxies may experience ram-pressure stripping. Both of these mechanisms are known to be associated with nuclear activity in galaxies (e.g. Sanders et al. 1988; Poggianti et al. 2017). Ergo, whilst we control for cluster membership, not controlling for position within the cluster may bias this analysis, i.e. if a central LERG is compared to a satellite control galaxy.

Using the brightest cluster galaxy (BCG) flag from the Yang et al. (2007) catalogue, we can crudely segregate galaxies into centrals and satellites. Doing this the galaxies in our sample selected as satellites all have projected radii from the BCG of $> 0.5R_{180}$, with 85% of these at $R > R_{180}$. When limiting the control sample to just those with the same central/satellite classification as the LERGs, the deficit of minor merger in LERGs is still observed at $> 3.5\sigma$ significance. This is dominated by the satellite galaxy population, where the of LERG minor merger deficit is seen at the 3σ level. On the other hand, in BCGs the minor merger fractions of LERGs and control galaxies are consistent, with $25.0 \pm 12.5\%$ of BCG LERGs and $30.3 \pm 9.1\%$ of BCG controls observed to be experiencing such interactions.

This suggests that minor mergers in the cluster core do not inhibit LERG activity. However, we note that just twelve galaxies in our sample, five of which are LERGs, are BCGs. Indeed only 18 LERGs in total within our sample lie within halos of $M_{180} > 10^{14} M_{\odot}$. That is to say, $28 \pm 11\%$ of our LERGs in clusters are BCGs. Given the frequent association of LERGs with BCGs (e.g. Tremblay et al. 2018), this might seem to be a relatively low number. To check that such a low BCG fraction amongst our cluster LERG sample should be expected, we cross match the entire Best & Heckman (2012) catalogue of LERGs at $z < 0.07$ (605 galaxies) with the Yang et al. (2007) group catalogue. Of the 42 LERGs found to be in clusters, $33 \pm 7\%$ are flagged as BCGs, consistent with what we observe in our data. Repeating this test with $z < 0.2$ shows $\sim 50\%$ of cluster LERGs to be BCGs, indicating that the low fraction of our cluster LERGs that are BCGs may be an effect of the low-redshift, $z < 0.07$, nature of this work. Indeed, Ching et al. (2017) demonstrate that low-power, $L_{1.4\text{GHz}} < 10^{24} \text{ W Hz}^{-1}$, radio AGN are not significantly more likely to be central galaxies than their radio quiet counterparts once matched on stellar mass and colour. A consequence of the limited depth of our sample is that

the median radio luminosity of our LERGs is an order of magnitude lower than this (see Figure 1). Therefore one might not expect our sample of LERGs to be hosted predominantly by BCGs. Furthermore, no account is taken of the structure of the clusters, in terms of whether it is relaxed, or actively coalescing with another structure (e.g., as is the case for Abell 1882, see Owers et al. 2013). Consequently we would urge further studies, making use of even deeper imaging, possibly from future facilities such as, e.g., the Large Synoptic Survey Telescope (LSST Science Collaboration et al. 2009), to obtain a larger volume of cluster LERGs with which to analyse merger trends.

6. SUMMARY

In this work we have exploited deep optical imaging to test the role of mergers in the evolution of inefficiently accreting RLAGN. This was achieved by comparing the prevalence and intensity of tidal features in LERGs (282 galaxies) and a control sample (1,622 galaxies) matched on redshift, galactic stellar mass and environment. In particular the depth of imaging, $\mu_r < 28 \text{ mag arcsec}^{-2}$, allowed for a large scale analysis of the role of minor mergers, a hypothesised LERG trigger (e.g., Kaviraj 2014a; Ellison et al. 2015), in LERG fuelling for the first time. Our main observations are:

1. No excess of minor mergers is observed in the LERG population relative to the control in any mass regime or large-scale environment. This is at odds with the hypothesis that minor mergers may present a fuel supply with which to power a weakly accreting RLAGN.
2. LERGs in clusters have a minor merger fraction of $8.8 \pm 2.9\%$ in contrast to $23.0 \pm 2.0\%$ of control galaxies, a $> 4\sigma$ deficit. This observation is not only inconsistent with the hypothesis that such events are a major contributor to LERG activity, but also suggests that minor mergers in the cluster environment act to prevent the evolution of an inactive galaxy into a LERG.
3. A significant, $> 4\sigma$, excess of major mergers is observed in relatively low-mass LERGs. At $M_* \lesssim 10^{11} M_{\odot}$, $10 \pm 1.5\%$ of these AGN are experiencing such large-scale interactions compared to $3.2 \pm 0.4\%$ of the control population. At higher masses the LERG major merger fraction tends towards that of the control population, with no LERG excess observed at $M_* > 10^{11.3} M_{\odot}$. This effect is seen most strongly for field galaxies, but we note that in all environments, any excess of

LERG major mergers is seen with the highest confidence at lower stellar masses.

In conclusion, our observations show that minor mergers do not fuel LERGs, and are in agreement with an overall picture where the majority of traditional LERGs are powered by the accretion of matter from the halo. A minority of lower mass LERGs are clearly associated with major mergers. In these cases we hypothesize that we may be witnessing a relatively brief phase of low-excitation in a galaxy that may evolve into a HERG.

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REFERENCES

- Abazajian, K. N., Adelman-McCarthy, J. K., Agüeros, M. A., et al. 2009, *ApJS*, 182, 543, doi: [10.1088/0067-0049/182/2/543](https://doi.org/10.1088/0067-0049/182/2/543)
- Allen, S. W., Dunn, R. J. H., Fabian, A. C., Taylor, G. B., & Reynolds, C. S. 2006, *MNRAS*, 372, 21, doi: [10.1111/j.1365-2966.2006.10778.x](https://doi.org/10.1111/j.1365-2966.2006.10778.x)
- Astropy Collaboration, Robitaille, T. P., Tollerud, E. J., et al. 2013, *A&A*, 558, A33, doi: [10.1051/0004-6361/201322068](https://doi.org/10.1051/0004-6361/201322068)
- Astropy Collaboration, Price-Whelan, A. M., Sipcz, B. M., et al. 2018, *AJ*, 156, 123, doi: [10.3847/1538-3881/aabc4f](https://doi.org/10.3847/1538-3881/aabc4f)
- Baldwin, J. A., Phillips, M. M., & Terlevich, R. 1981, *PASP*, 93, 5, doi: [10.1086/130766](https://doi.org/10.1086/130766)
- Barsanti, S., Owers, M. S., Brough, S., et al. 2018, *ApJ*, 857, 71, doi: [10.3847/1538-4357/aab61a](https://doi.org/10.3847/1538-4357/aab61a)
- Baum, S. A., Heckman, T. M., & van Breugel, W. 1992, *ApJ*, 389, 208, doi: [10.1086/171198](https://doi.org/10.1086/171198)
- Baum, S. A., Zirbel, E. L., & O’Dea, C. P. 1995, *ApJ*, 451, 88, doi: [10.1086/176202](https://doi.org/10.1086/176202)
- Becker, R. H., White, R. L., & Helfand, D. J. 1995, *ApJ*, 450, 559, doi: [10.1086/176166](https://doi.org/10.1086/176166)
- Bessiere, P. S., Tadhunter, C. N., Ramos Almeida, C., & Villar Martín, M. 2014, *MNRAS*, 438, 1839, doi: [10.1093/mnras/stt2333](https://doi.org/10.1093/mnras/stt2333)
- Best, P. N., & Heckman, T. M. 2012, *MNRAS*, 421, 1569, doi: [10.1111/j.1365-2966.2012.20414.x](https://doi.org/10.1111/j.1365-2966.2012.20414.x)
- Best, P. N., Kauffmann, G., Heckman, T. M., et al. 2005a, *MNRAS*, 362, 25, doi: [10.1111/j.1365-2966.2005.09192.x](https://doi.org/10.1111/j.1365-2966.2005.09192.x)
- Best, P. N., Kauffmann, G., Heckman, T. M., & Ivezić, Ž. 2005b, *MNRAS*, 362, 9, doi: [10.1111/j.1365-2966.2005.09283.x](https://doi.org/10.1111/j.1365-2966.2005.09283.x)
- Best, P. N., Von Der Linden, A., Kauffmann, G., Heckman, T. M., & Kaiser, C. R. 2007, *MNRAS*, 379, 894, doi: [10.1111/j.1365-2966.2007.11937.x](https://doi.org/10.1111/j.1365-2966.2007.11937.x)
- Blandford, R. D., & Znajek, R. L. 1977, *MNRAS*, 179, 433, doi: [10.1093/mnras/179.3.433](https://doi.org/10.1093/mnras/179.3.433)
- Blanton, M. R., Lin, H., Lupton, R. H., et al. 2003, *AJ*, 125, 2276, doi: [10.1086/344761](https://doi.org/10.1086/344761)
- Blum, R. D., Burleigh, K., Dey, A., et al. 2016, *Am. Astron. Soc. Meet. Abstr.*, 228, 317.01
- Bournaud, F., Dekel, A., Teyssier, R., et al. 2011, *ApJ*, 741, L33, doi: [10.1088/2041-8205/741/2/L33](https://doi.org/10.1088/2041-8205/741/2/L33)
- Buttiglione, S., Capetti, A., Celotti, A., et al. 2010, *A&A*, 509, 6, doi: [10.1051/0004-6361/200913290](https://doi.org/10.1051/0004-6361/200913290)
- Cameron, E. 2011, *PASA*, 28, 128, doi: [10.1071/AS10046](https://doi.org/10.1071/AS10046)
- Carpineti, A., Kaviraj, S., Darg, D., et al. 2012, *MNRAS*, 420, 2139, doi: [10.1111/j.1365-2966.2011.20179.x](https://doi.org/10.1111/j.1365-2966.2011.20179.x)
- Chiaberge, M., Gilli, R., M. Lotz, J., & Norman, C. 2015, *ApJ*, 806, 147, doi: [10.1088/0004-637X/806/2/147](https://doi.org/10.1088/0004-637X/806/2/147)
- Ching, J. H. Y., Croom, S. M., Sadler, E. M., et al. 2017, *MNRAS*, 469, 4584, doi: [10.1093/mnras/stx1173](https://doi.org/10.1093/mnras/stx1173)
- Comerford, J. M., Barrows, R. S., Müller-Sánchez, F., et al. 2017, *ApJ*, 849, 102, doi: [10.3847/1538-4357/aa8e4b](https://doi.org/10.3847/1538-4357/aa8e4b)
- Condon, J. J., Cotton, W. D., Greisen, E. W., et al. 1998, *AJ*, 115, 1693, doi: [10.1086/300337](https://doi.org/10.1086/300337)
- Davies, R. I., Hicks, E., Erwin, P., et al. 2017, *MNRAS*, 466, 4917, doi: [10.1093/mnras/stx045](https://doi.org/10.1093/mnras/stx045)
- de Jong, J. T. A., Verdoes Kleijn, G. A., Kuijken, K. H., & Valentijn, E. A. 2013, *Exp. Astron.*, 35, 25, doi: [10.1007/s10686-012-9306-1](https://doi.org/10.1007/s10686-012-9306-1)
- De Rijcke, S., Prugniel, P., Simien, F., & Dejonghe, H. 2006, *MNRAS*, 369, 1321, doi: [10.1111/j.1365-2966.2006.10377.x](https://doi.org/10.1111/j.1365-2966.2006.10377.x)
- Donahue, M., de Messières, G. E., O’Connell, R. W., et al. 2011, *ApJ*, 732, 40, doi: [10.1088/0004-637X/732/1/40](https://doi.org/10.1088/0004-637X/732/1/40)
- Driver, S. P., Wright, A. H., Andrews, S. K., et al. 2016, *MNRAS*, 455, 3911, doi: [10.1093/mnras/stv2505](https://doi.org/10.1093/mnras/stv2505)
- Duc, P.-A., Cuillandre, J.-C., Karabal, E., et al. 2015, *MNRAS*, 446, 120, doi: [10.1093/mnras/stu2019](https://doi.org/10.1093/mnras/stu2019)
- Edge, A. 2001, *MNRAS*, 328, 762, doi: [10.1046/j.1365-8711.2001.04802.x](https://doi.org/10.1046/j.1365-8711.2001.04802.x)
- Ellison, S. L., Patton, D. R., & Hickox, R. C. 2015, *MNRAS*, 451, 35, doi: [10.1093/mnrasl/slv061](https://doi.org/10.1093/mnrasl/slv061)
- Evans, D. A., Summers, A. C., Hardcastle, M. J., et al. 2011, *ApJ*, 741, L4, doi: [10.1088/2041-8205/741/1/L4](https://doi.org/10.1088/2041-8205/741/1/L4)
- Fabian, A., & Rees, M. 1995, *MNRAS*, 277, L55, doi: [10.1093/mnras/277.1.L55](https://doi.org/10.1093/mnras/277.1.L55)
- Fliri, J., & Trujillo, I. 2016, *MNRAS*, 456, 1359, doi: [10.1093/mnras/stv2686](https://doi.org/10.1093/mnras/stv2686)
- Gaspari, M., Ruszkowski, M., & Oh, S. P. 2013, *MNRAS*, 432, 3401, doi: [10.1093/mnras/stt692](https://doi.org/10.1093/mnras/stt692)
- Gaspari, M., Temi, P., & Brighenti, F. 2017, *MNRAS*, 466, 677, doi: [10.1093/mnras/stw3108](https://doi.org/10.1093/mnras/stw3108)
- Gatti, M., Shankar, F., Bouillot, V., et al. 2016, *MNRAS*, 456, 1073, doi: [10.1093/mnras/stv2754](https://doi.org/10.1093/mnras/stv2754)
- Gordon, Y. A., Pimblet, K. A., & Owers, M. S. 2018a, *Res. Notes AAS*, 2, 132, doi: [10.3847/2515-5172/aad596](https://doi.org/10.3847/2515-5172/aad596)
- Gordon, Y. A., Owers, M. S., Pimblet, K. A., et al. 2017, *MNRAS*, 465, 2671, doi: [10.1093/mnras/stw2925](https://doi.org/10.1093/mnras/stw2925)
- Gordon, Y. A., Pimblet, K. A., Owers, M. S., et al. 2018b, *MNRAS*, 475, 4223, doi: [10.1093/mnras/sty115](https://doi.org/10.1093/mnras/sty115)
- Haines, C. P., Pereira, M. J., Sanderson, A. J. R., et al. 2012, *ApJ*, 754, 97, doi: [10.1088/0004-637X/754/2/97](https://doi.org/10.1088/0004-637X/754/2/97)
- Hardcastle, M. J., Evans, D. A., & Croston, J. H. 2006, *MNRAS*, 370, 1893, doi: [10.1111/j.1365-2966.2006.10615.x](https://doi.org/10.1111/j.1365-2966.2006.10615.x)
- . 2007, *MNRAS*, 376, 1849, doi: [10.1111/j.1365-2966.2007.11572.x](https://doi.org/10.1111/j.1365-2966.2007.11572.x)

- Heckman, T. M., Ptak, A., Hornschemeier, A., & Kauffmann, G. 2005, *ApJ*, 634, 161, doi: [10.1086/491665](https://doi.org/10.1086/491665)
- Heckman, T. M., Smith, E. P., Baum, S. A., et al. 1986, *ApJ*, 311, 526, doi: [10.1086/164793](https://doi.org/10.1086/164793)
- Hill, G. J., & Lilly, S. J. 1991, *ApJ*, 367, 1, doi: [10.1086/169597](https://doi.org/10.1086/169597)
- Hood, C. E., Kannappan, S. J., Stark, D. V., et al. 2018, *ApJ*, 857, 144, doi: [10.3847/1538-4357/aab719](https://doi.org/10.3847/1538-4357/aab719)
- Jackson, N., & Rawlings, S. 1997, *MNRAS*, 286, 241, doi: [10.1093/mnras/286.1.241](https://doi.org/10.1093/mnras/286.1.241)
- Ji, L., Peirani, S., & Yi, S. K. 2014, *A&A*, 566, A97, doi: [10.1051/0004-6361/201423530](https://doi.org/10.1051/0004-6361/201423530)
- Kado-Fong, E., Greene, J. E., Hendel, D., et al. 2018, *ApJ*, 866, 103, doi: [10.3847/1538-4357/aae0f0](https://doi.org/10.3847/1538-4357/aae0f0)
- Kauffmann, G., Heckman, T. M., & Best, P. N. 2008, *MNRAS*, 384, 953, doi: [10.1111/j.1365-2966.2007.12752.x](https://doi.org/10.1111/j.1365-2966.2007.12752.x)
- Kauffmann, G., Heckman, T. M., Tremonti, C., et al. 2003a, *MNRAS*, 346, 1055, doi: [10.1111/j.1365-2966.2003.07154.x](https://doi.org/10.1111/j.1365-2966.2003.07154.x)
- Kauffmann, G., Heckman, T. M., Simon White, D. M., et al. 2003b, *MNRAS*, 341, 33, doi: [10.1046/j.1365-8711.2003.06291.x](https://doi.org/10.1046/j.1365-8711.2003.06291.x)
- Kaviraj, S. 2010, *MNRAS*, 406, 382, doi: [10.1111/j.1365-2966.2010.16714.x](https://doi.org/10.1111/j.1365-2966.2010.16714.x)
- . 2014a, *MNRAS*, 440, 2944, doi: [10.1093/mnras/stu338](https://doi.org/10.1093/mnras/stu338)
- . 2014b, *MNRAS*, 437, L41, doi: [10.1093/mnrasl/slt136](https://doi.org/10.1093/mnrasl/slt136)
- Kaviraj, S., Laigle, C., Kimm, T., et al. 2017, *MNRAS*, 467, 4739, doi: [10.1093/mnras/stx126](https://doi.org/10.1093/mnras/stx126)
- King, A., & Nixon, C. 2015, *MNRAS*, 453, L46, doi: [10.1093/mnrasl/slv098](https://doi.org/10.1093/mnrasl/slv098)
- . 2018, *ApJ*, 857, L7, doi: [10.3847/2041-8213/aab8f9](https://doi.org/10.3847/2041-8213/aab8f9)
- Lintott, C. J., Schawinski, K., Slosar, A., et al. 2008, *MNRAS*, 389, 1179, doi: [10.1111/j.1365-2966.2008.13689.x](https://doi.org/10.1111/j.1365-2966.2008.13689.x)
- Lofthouse, E. K., Kaviraj, S., Smith, D. J. B., & Hardcastle, M. J. 2018, *MNRAS*, 479, 807, doi: [10.1093/mnras/sty1375](https://doi.org/10.1093/mnras/sty1375)
- Lotz, J. M., Jonsson, P., Cox, T. J., & Primack, J. R. 2008, *MNRAS*, 391, 1137, doi: [10.1111/j.1365-2966.2008.14004.x](https://doi.org/10.1111/j.1365-2966.2008.14004.x)
- . 2010, *MNRAS*, 404, 575, doi: [10.1111/j.1365-2966.2010.16268.x](https://doi.org/10.1111/j.1365-2966.2010.16268.x)
- Loveday, J. 1997, *ApJ*, 489, 29, doi: [10.1086/304778](https://doi.org/10.1086/304778)
- LSST Science Collaboration, Abell, P. A., Allison, J., et al. 2009, arXiv e-prints, 0912.0201. <https://arxiv.org/abs/0912.0201>
- Lupton, R., Blanton, M. R., Fekete, G., et al. 2004, *PASP*, 116, 133, doi: [10.1086/382245](https://doi.org/10.1086/382245)
- Martin, G., Kaviraj, S., Volonteri, M., et al. 2018, *MNRAS*, 476, 2801, doi: [10.1093/mnras/sty324](https://doi.org/10.1093/mnras/sty324)
- Mingo, B., Hardcastle, M. J., Croston, J. H., et al. 2014, *MNRAS*, 440, 269, doi: [10.1093/mnras/stu263](https://doi.org/10.1093/mnras/stu263)
- Morales, G., Martínez-Delgado, D., Grebel, E. K., et al. 2018, *A&A*, 614, A143, doi: [10.1051/0004-6361/201732271](https://doi.org/10.1051/0004-6361/201732271)
- Narayan, R., & Yi, I. 1995, *ApJ*, 444, 231, doi: [10.1086/175599](https://doi.org/10.1086/175599)
- Nealon, R., Price, D. J., & Nixon, C. J. 2015, *MNRAS*, 448, 1526, doi: [10.1093/mnras/stv014](https://doi.org/10.1093/mnras/stv014)
- O’Dea, C. P., Baum, S. A., Maloney, P. R., Tacconi, L. J., & Sparks, W. B. 1994, *ApJ*, 422, 467, doi: [10.1086/173742](https://doi.org/10.1086/173742)
- O’Dea, C. P., Baum, S. A., Privon, G., et al. 2008, *ApJ*, 681, 1035, doi: [10.1086/588212](https://doi.org/10.1086/588212)
- Owers, M. S., Baldry, I. K., Bauer, A. E., et al. 2013, *ApJ*, 772, 104, doi: [10.1088/0004-637X/772/2/104](https://doi.org/10.1088/0004-637X/772/2/104)
- Pace, C., & Salim, S. 2014, *ApJ*, 785, 66, doi: [10.1088/0004-637X/785/1/66](https://doi.org/10.1088/0004-637X/785/1/66)
- Patton, D. R., & Atfield, J. E. 2008, *ApJ*, 685, 235, doi: [10.1086/590542](https://doi.org/10.1086/590542)
- Pimblet, K. A., Shabala, S. S., Haines, C. P., Fraser-McKelvie, A., & Floyd, D. J. E. 2013, *MNRAS*, 429, 1827, doi: [10.1093/mnras/sts470](https://doi.org/10.1093/mnras/sts470)
- Pipino, A., Kaviraj, S., Bildfell, C., et al. 2009, *MNRAS*, 395, 462, doi: [10.1111/j.1365-2966.2009.14534.x](https://doi.org/10.1111/j.1365-2966.2009.14534.x)
- Poggianti, B. M., Jaffé, Y. L., Moretti, A., et al. 2017, *Nature*, 548, 304, doi: [10.1038/nature23462](https://doi.org/10.1038/nature23462)
- Pop, A.-R., Pillepich, A., Amorisco, N. C., & Hernquist, L. 2018, *MNRAS*, 480, 1715, doi: [10.1093/mnras/sty1932](https://doi.org/10.1093/mnras/sty1932)
- Ramos Almeida, C., Bessiere, P. S., Tadhunter, C. N., et al. 2013, *MNRAS*, 436, 997, doi: [10.1093/mnras/stt1595](https://doi.org/10.1093/mnras/stt1595)
- . 2012, *MNRAS*, 419, 687, doi: [10.1111/j.1365-2966.2011.19731.x](https://doi.org/10.1111/j.1365-2966.2011.19731.x)
- Richards, G. T., Fan, X., Newberg, H. J., et al. 2002, *AJ*, 123, 2945, doi: [10.1086/340187](https://doi.org/10.1086/340187)
- Robotham, A., Driver, S. P., Norberg, P., et al. 2010, *PASA*, 27, 76, doi: [10.1071/AS09053](https://doi.org/10.1071/AS09053)
- Sabater, J., Best, P. N., & Argudo-Fernández, M. 2013, *MNRAS*, 430, 638, doi: [10.1093/mnras/sts675](https://doi.org/10.1093/mnras/sts675)
- Salpeter, E. E. 1964, *ApJ*, 140, 796, doi: [10.1086/147973](https://doi.org/10.1086/147973)
- Sanders, D. B., Soifer, B. T., Elias, J. H., et al. 1988, *ApJ*, 325, 74, doi: [10.1086/165983](https://doi.org/10.1086/165983)
- Schawinski, K., Koss, M., Berney, S., & Sartori, L. F. 2015, *MNRAS*, 451, 2517, doi: [10.1093/mnras/stv1136](https://doi.org/10.1093/mnras/stv1136)
- Scott, C., & Kaviraj, S. 2014, *MNRAS*, 437, 2137, doi: [10.1093/mnras/stt2014](https://doi.org/10.1093/mnras/stt2014)
- Shabala, S. S., Deller, A., Kaviraj, S., et al. 2017, *MNRAS*, 464, 4706, doi: [10.1093/mnras/stw2536](https://doi.org/10.1093/mnras/stw2536)
- Smith, R. J., Hudson, M. J., Nelan, J. E., et al. 2004, *AJ*, 128, 1558, doi: [10.1086/423915](https://doi.org/10.1086/423915)

- Smolčić, V. 2009, *ApJ*, 699, 43,
doi: [10.1088/0004-637X/699/1/L43](https://doi.org/10.1088/0004-637X/699/1/L43)
- Strauss, M. A., Weinberg, D. H., Lupton, R. H., & Narayanan, V. K. 2002, *AJ*, 124, 1810,
doi: [10.1086/342343](https://doi.org/10.1086/342343)
- Tadhunter, C., Dickson, R., Morganti, R., et al. 2002, *MNRAS*, 330, 977, doi: [10.1046/j.1365-8711.2002.05153.x](https://doi.org/10.1046/j.1365-8711.2002.05153.x)
- Tadhunter, C., Robinson, T. G., González Delgado, R. M., Wills, K., & Morganti, R. 2005, *MNRAS*, 356, 480,
doi: [10.1111/j.1365-2966.2004.08470.x](https://doi.org/10.1111/j.1365-2966.2004.08470.x)
- Tadhunter, C. N., Morganti, R., Robinson, A., et al. 1998, *MNRAS*, 298, 1035,
doi: [10.1046/j.1365-8711.1998.01706.x](https://doi.org/10.1046/j.1365-8711.1998.01706.x)
- Tremblay, G. R., Oonk, J. B. R., Combes, F., et al. 2016, *Nature*, 534, 218, doi: [10.1038/nature17969](https://doi.org/10.1038/nature17969)
- Tremblay, G. R., Combes, F., Oonk, J. B. R., et al. 2018, *ApJ*, 865, 13, doi: [10.3847/1538-4357/aad6dd](https://doi.org/10.3847/1538-4357/aad6dd)
- Tremonti, C. A., Heckman, T. M., Kauffmann, G., et al. 2004, *ApJ*, 613, 898, doi: [10.1086/423264](https://doi.org/10.1086/423264)
- Villforth, C., Hamilton, T., Pawlik, M. M., et al. 2017, *MNRAS*, 466, 812, doi: [10.1093/mnras/stw3037](https://doi.org/10.1093/mnras/stw3037)
- Voit, G. M., Meece, G., Li, Y., et al. 2017, *ApJ*, 845, 80,
doi: [10.3847/1538-4357/aa7d04](https://doi.org/10.3847/1538-4357/aa7d04)
- Weston, M. E., McIntosh, D. H., Brodwin, M., et al. 2017, *MNRAS*, 464, 3882, doi: [10.1093/mnras/stw2620](https://doi.org/10.1093/mnras/stw2620)
- Willett, K. W., Lintott, C. J., Bamford, S. P., et al. 2013, *MNRAS*, 435, 2835, doi: [10.1093/mnras/stt1458](https://doi.org/10.1093/mnras/stt1458)
- Yang, X., Mo, H. J., van den Bosch, F. C., et al. 2007, *ApJ*, 671, 153, doi: [10.1086/522027](https://doi.org/10.1086/522027)
- Yoon, J. H., Schawinski, K., Sheen, Y.-K., Ree, C. H., & Yi, S. K. 2008, *ApJS*, 176, 414, doi: [10.1086/528958](https://doi.org/10.1086/528958)
- York, D. G., Adelman, J., Anderson, J. E., et al. 2000, *AJ*, 120, 1579, doi: [10.1086/301513](https://doi.org/10.1086/301513)
- Zirbel, E. L. 1997, *ApJ*, 476, 489, doi: [10.1086/303626](https://doi.org/10.1086/303626)