



<b>Title</b>	<b>Formation of Dense Molecular Gas and Stars at the Circumnuclear Starburst Ring in the Barred Galaxy NGC 7552</b>
<b>Author(s)</b>	<b>Pan, H A; Lim, J J L; Matsushita, S; Wong, T; Ryder, S</b>
<b>Citation</b>	<b>The Astrophysical Journal, 2013, v. 768, p. 57</b>
<b>Issued Date</b>	<b>2013</b>
<b>URL</b>	<b><a href="http://hdl.handle.net/10722/186170">http://hdl.handle.net/10722/186170</a></b>
<b>Rights</b>	<b>Creative Commons: Attribution 3.0 Hong Kong License</b>

## FORMATION OF DENSE MOLECULAR GAS AND STARS AT THE CIRCUMNUCLEAR STARBURST RING IN THE BARRED GALAXY NGC 7552

HSI-AN PAN<sup>1,2,3</sup>, JEREMY LIM<sup>4,7</sup>, SATOKI MATSUSHITA<sup>3</sup>, TONY WONG<sup>5</sup>, AND STUART RYDER<sup>6</sup>

<sup>1</sup> Department of Astronomical Science, The Graduate University for Advanced Studies, Shonan Village, Hayama,  
Kanagawa 240-0193, Japan; [pan.h.a@nao.ac.jp](mailto:pan.h.a@nao.ac.jp)

<sup>2</sup> Nobeyama Radio Observatory, NAOJ, Minamimaki, Minamisaku, Nagano 384-1305, Japan

<sup>3</sup> Academia Sinica, Institute of Astronomy & Astrophysics (ASIAA), P.O. Box 23-141, Taipei 10617, Taiwan

<sup>4</sup> Department of Physics, University of Hong Kong, Pokfulam, Hong Kong

<sup>5</sup> Department of Astronomy, University of Illinois, 1002 W. Green Street, Urbana, IL 61801, USA

<sup>6</sup> Australian Astronomical Observatory, P.O. Box 915, North Ryde, NSW 1670, Australia

Received 2012 December 24; accepted 2013 March 10; published 2013 April 15

### ABSTRACT

We present millimeter molecular line complemented by optical observations, along with a reanalysis of archival centimeter HI and continuum data, to infer the global dynamics and determine where dense molecular gas and massive stars preferentially form in the circumnuclear starburst ring of the barred-spiral galaxy NGC 7552. We find diffuse molecular gas in a pair of dust lanes each running along the large-scale galactic bar, as well as in the circumnuclear starburst ring. We do not detect dense molecular gas in the dust lanes, but find such gas concentrated in two knots where the dust lanes make contact with the circumnuclear starburst ring. When convolved to the same angular resolution as the images in dense gas, the radio continuum emission of the circumnuclear starburst ring also exhibits two knots, each lying downstream of an adjacent knot in dense gas. The results agree qualitatively with the idea that massive stars form from dense gas at the contact points, where diffuse gas is channeled into the ring along the dust lanes, and later explode as supernovae downstream of the contact points. Based on the inferred rotation curve, however, the propagation time between the respective pairs of dense gas and centimeter continuum knots is about an order of magnitude shorter than the lifetimes of OB stars. We discuss possible reasons for this discrepancy, and conclude that either the initial mass function is top-heavy or massive stars in the ring do not form exclusively at the contact points where dense molecular gas is concentrated.

*Key words:* galaxies: individual (NGC 7552) – galaxies: ISM – galaxies: kinematics and dynamics – galaxies: starburst

*Online-only material:* color figures

### 1. INTRODUCTION

Over half (e.g.,  $\gtrsim 70\%$ ; Eskridge et al. 2000) of all spiral galaxies in the local universe exhibit large-scale bars. Such barred galaxies exhibit circumnuclear starbursts more frequently than non-barred spiral galaxies (Jogee et al. 2005; Ellison et al. 2011). The circumnuclear starburst usually takes the form of a ring centered on the nucleus (e.g., review by Buta & Combes 1996), and has a star formation rate (SFR) that can make up a significant fraction, if not dominate, the total SFR of the galaxy (Jogee et al. 2005). Is gas transported from the surrounding regions toward the center of the galaxy to fuel the circumnuclear starburst, and if so how? If bars transport gas inward, then why we do not see circumnuclear starbursts in virtually all barred spiral galaxies? Stated in a different way, why is star formation is strongly promoted in the circumnuclear region of some galaxies? In this work, we address the last question through observations of the barred spiral galaxy NGC 7552, which exhibits a circumnuclear starburst ring with an SFR of  $10\text{--}15 M_{\odot} \text{ yr}^{-1}$  that comprises at least half of the total SFR of this galaxy.

Theoretical studies suggest that, in galactic disks, gravitational torques induced by bar potentials drive gas inflows. In this scenario, molecular gas experiences two principal shocks:

(1) at the two leading edges of the large-scale bar, where dust lanes (indicating, presumably, the pileup of molecular gas) are commonly seen (e.g., Sheth et al. 2002); and (2) at the two locations where the inner terminal portions of the bar make contact with the circumnuclear starburst ring, where twin peaks in molecular gas as traced by CO are commonly observed (e.g., Kenney et al. 1992; Sakamoto et al. 1999). During the passage of molecular gas in the galactic disk through the bar, angular momentum is transferred from the gas to the bar through the aforementioned shocks, resulting in a net inflow of gas along the dust lanes. The inflowing gas is eventually halted at an inner Lindblad resonance (ILR; or, if there are two ILRs, then between the inner ILR and outer ILR), which has the form of a ring, where the bar terminates and within which the net torque exerted by the bar is zero (Buta & Combes 1996). Because the formation of a circumnuclear gas ring requires a bar that is strong enough (i.e., to induce a sufficiently strong non-axisymmetric potential) to promote adequate gas inflow, the mere presence of a bar is not sufficient for the presence of a circumnuclear gas ring. Simulations (e.g., Regan & Teuben 2004), as well as observations (e.g., Sakamoto et al. 1999; Laurikainen et al. 2007), suggest that galaxies with long and massive (morphological classification SBa) bars are more effective at channeling gas toward their centers than those with weak bars (SBc), and hence preferentially host circumnuclear gas rings.

Although molecular gas has been observed to accumulate along the leading edges of the bar and at the ILR in agreement

<sup>7</sup> Adjunct Research Fellow, Academia Sinica, Institute of Astronomy, and Astrophysics (ASIAA), P.O. Box 23-141, Taipei 10617, Taiwan.

with model predictions, there is a striking difference between the observed properties of the molecular gas and the star formation activity in these two regions. The molecular gas in (at the leading edge of) the bar has been found to be much more diffuse on average (e.g., Downes et al. 1996; Hüttemeister et al. 2000; Schinnerer et al. 2002), and the star formation activity in the bar also much less pronounced, than that in the circumnuclear starburst ring. Based on observations of NGC 6951, a strongly barred spiral galaxy that exhibits a circumnuclear starburst ring and which in many other respects are remarkably similar to NGC 7552 as explained below, Kohno et al. (1999) described how diffuse molecular gas may collect in the circumnuclear starburst ring of this galaxy and the circumstances under which this gas subsequently becomes sufficiently dense to form stars. They found that, just like in a number of other barred galaxies having circumnuclear starburst rings, diffuse molecular gas as traced in  $^{12}\text{CO}$  is concentrated in two knots on opposite sides of the ring where a pair of dust lanes, each running along the (leading edge of the) bar on opposite sides of the nucleus, make contact with the ring, demonstrating that the twin CO peaks likely represent molecular gas channeled inward along the dust lanes. For the first time, they showed that dense molecular gas as traced in HCN is not cospatial with the CO knots, but is instead concentrated in two knots located just downstream of the respective CO knots. Kohno et al. (1999) argued that diffuse molecular gas, channeled inward along  $x_1$  orbits (closely traced by the dust lanes), is too turbulent to become dense where they intersect with  $x_2$  orbits (traced by the circumnuclear starburst ring), as the gas at these locations (delineated by the CO knots) experience relatively strong shocks. Instead, they proposed that the molecular gas only becomes dense further downstream where the aforementioned turbulence has partially dissipated, leading to the collapse of gas clouds via gravitational (Toomre-type) instabilities (Toomre 1964) and subsequent star formation.

Because intercomparisons can provide insights not possible from studies of a single object alone, here we test the idea proposed by Kohno et al. (1999) on another galaxy. NGC 7552 is a member of the Southern Grus Quartet at a distance of 22.3 Mpc (so that  $1'' = 107$  pc; Kennicutt et al. 2003). Observations of the entire group in atomic hydrogen (H I) gas have revealed an extension on the southwestern side of NGC 7552 indicative of tidal interactions with one or more group members (Dahlem 2005), although no such distortion is visible in the optical disk suggesting that only the outskirts of NGC 7552 is strongly perturbed. The gross optical morphology of NGC 7552—circumnuclear starburst ring, large-scale bar with dust lanes along the leading edges, and two prominent spiral arms emerging from the ends of the bar as can be seen in Figure 1—is remarkably similar to that of NGC 6951, including even the orientation of the abovementioned features in the sky. The large-scale stellar disk of NGC 7552 is close to face-on, with an axial ratio between the optical major and minor axes indicating an inclination of about  $25^\circ$  (for comparison, also based on axial ratios from optical images, the LEDA extragalactic database quotes an average inclination of  $23^\circ$ , and Feinstein et al. 1990 infer an inclination of  $28^\circ$ ). The nuclear activity of NGC 7552 is classified as a low-ionization nuclear emission-line region (LINER) with H II regions (Durret & Bergeron 1987). Claussen & Sahai (1992) find that the molecular gas in NGC 7552 as traced in  $^{12}\text{CO}$  (1–0) is strongly concentrated in the nucleus and the bar.

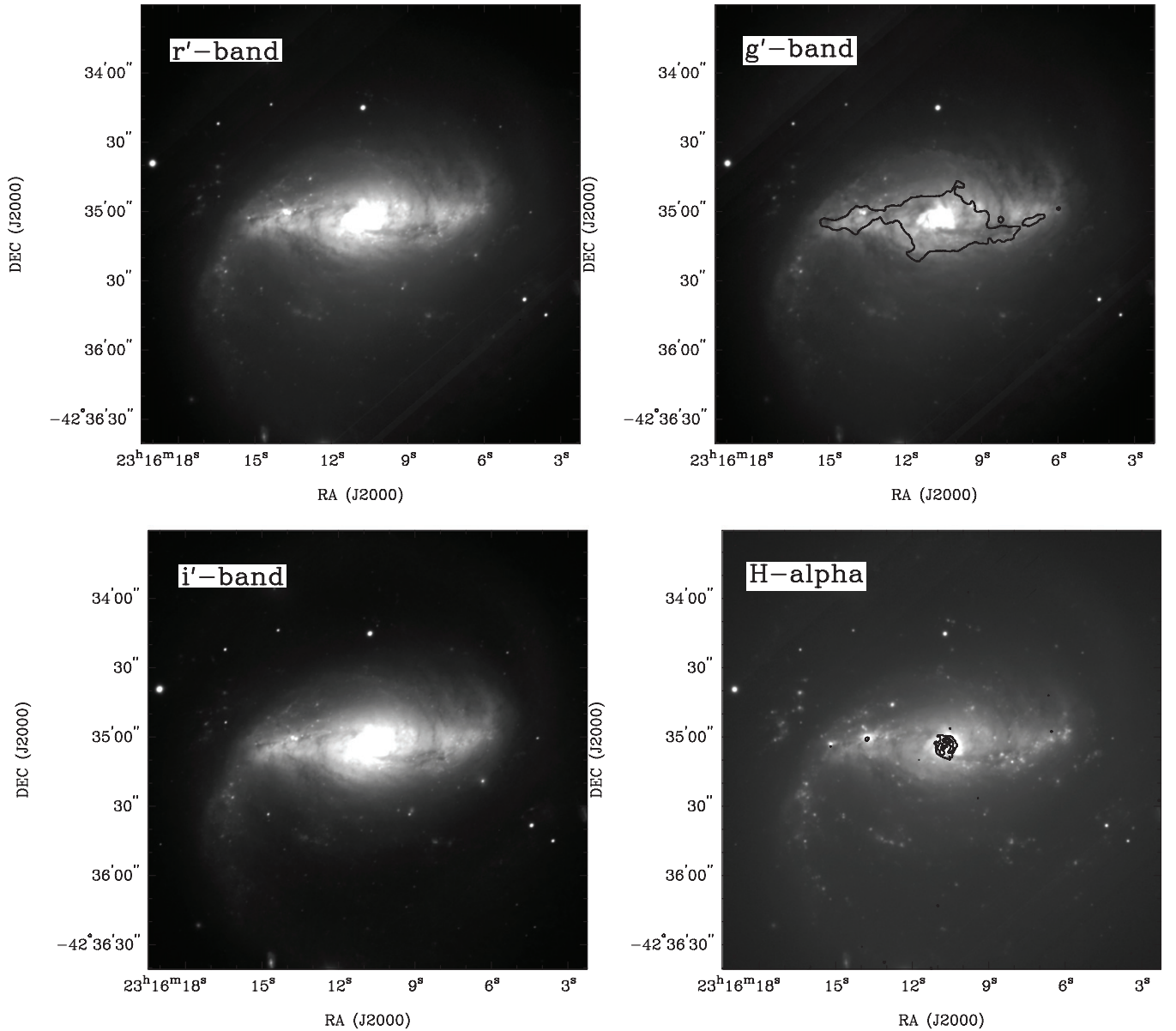
Forbes et al. (1994a) have imaged the circumnuclear starburst ring of NGC 7552 in the radio continuum at 6 cm and 3 cm,

providing the least-observed view of the ring so far available. As will be explained in greater detail later, the ring is nearly circular, and exhibits five prominent knots. The radio spectrum of each knot falls between 6 cm and 3 cm, indicating non-thermal synchrotron emission produced, presumably, as a result of supernova explosions. Forbes et al. (1994b) also imaged the ring in H $\alpha$  and Schinnerer et al. (1997) in Br $\gamma$ , where both these lines trace gas ionized by recently formed massive stars.

Our work presented here involves multi-line radio interferometric complemented by optical observations of NGC 7552, along with a reanalysis of archival data for this galaxy in H I and centimeter continuum. In Section 2 we provide details on the observations performed and the data reanalyzed, and in Section 3 the results obtained. In Section 4 we derive the physical properties of the molecular gas, and in Section 5 the global kinematics of NGC 7552. Then, in Section 6, we examine the implications of our results and analyses for the formation of dense molecular gas and stars in the circumnuclear starburst ring of NGC 7552, and propose an important revision to the picture advanced by Kohno et al. (1999), based on observations of the circumnuclear starburst ring in NGC 6951, that is better able to explain the observed properties of the circumnuclear starburst ring in NGC 7552 as well as more recent observations of dense gas in the circumnuclear starburst ring of NGC 6951. Finally, in Section 7, we provide a concise summary of our work.

## 2. OBSERVATIONS AND DATA REDUCTION

Our millimeter molecular-line observations toward the center of NGC 7552 focused on four bright tracers of molecular hydrogen gas,  $^{12}\text{CO}$ ,  $^{13}\text{CO}$ , HCN, and HCO $^+$ .  $^{12}\text{CO}$  traces relatively diffuse ( $\sim 10^2$ – $10^3$  cm $^{-3}$ ) molecular gas characteristic of the bulk of giant molecular clouds (GMCs) in our Galaxy, whereas HCN and HCO $^+$  traces relatively dense ( $\sim 10^5$ – $10^6$  cm $^{-3}$ ) molecular gas characteristic of dense condensations in GMCs where stars form; as we will show in Section 4.2, in our observations  $^{13}\text{CO}$  traces molecular gas at intermediate densities. The spatial distribution of  $^{12}\text{CO}$  therefore indicates locations where atomic hydrogen gas has presumably been compressed to form molecular hydrogen gas, whereas the spatial distribution of HCN and HCO $^+$  indicates locations where the molecular gas has attained densities sufficiently high to form stars. These molecular-line observations also provide information on the gas kinematics: as described in Section 5.1, we used the  $^{12}\text{CO}$  image to infer a rotation curve for the central region of NGC 7552. The H I image of NGC 7552 provides information on the gas kinematics beyond the central region imaged in molecular gas, and as described in Section 5.1 was used to infer a rotation curve for the galaxy between the innermost to the outermost detectable regions of the H I disk. The rotation curve thus inferred was used to derive the predicted locations of dynamical resonances in the galaxy, as well as to compute dynamical timescales in the circumnuclear starburst ring. We compared the predicted locations of dynamical resonances to individual features in NGC 7552 from our wide- and narrow-band optical observations, thus providing an assessment of the reliability of the inferred rotation curve. The narrow-band H $\alpha$  image from our optical observations also provided estimates of the SFRs in different parts of the galaxy. Finally, from the radio continuum image of the circumnuclear starburst ring, we inferred locations where large concentrations of massive stars recently met their demise as supernovae.



**Figure 1.** Optical  $r'$ -band (upper left),  $g'$ -band (upper right),  $i'$ -band (lower left), and  $H\alpha$  (lower right) images of NGC 7552 taken with GMOS on the Gemini South telescope. The *Spitzer*  $5.8\ \mu\text{m}$  contour at  $5.5\ \text{MJy sr}^{-1}$  (Kennicutt et al. 2003) is overlaid on the  $g'$ -band image, highlighting the orientation of the dust lanes. Likewise, the 3 cm continuum contours (Section 3.3; see text) showing the circumnuclear starburst ring are plotted in the  $H\alpha$  image, with contour steps of  $3\sigma$ ,  $15\sigma$ , and  $35\sigma$ , where  $1\sigma = 66\ \mu\text{Jy beam}^{-1}$ .

### 2.1. Optical Observations: Gemini South

The optical observations of NGC 7552 reported here are the product of a winning entry in the 2011 Australian Gemini School Astronomy Contest, for which one of us (S. Ryder) was a co-investigator on the contest proposal. Images of the galaxy in  $g'$ ,  $r'$ , and  $i'$  filters were obtained with the Gemini Multi-Object Spectrograph (GMOS; Hook et al. 2004) attached to the Gemini South Telescope as part of the program GS-2011A-Q-12 (PI: C. Onken). The filters are similar to the filters used by the Sloan Digital Sky Survey (SDSS). We also obtained an  $H\alpha$  image with a narrow-band  $H\alpha$  filter centered on 656 nm with a 7 nm wide bandpass. The observations were done in seeing between  $0''.6$ – $0''.8$  under photometric conditions on 2011 July 23 UT. Four exposures per filter were obtained with integration time between 100 and 300 s, with each exposure offset by  $10''$  spatially to allow filling-in of the inter-CCD gaps in GMOS. On-chip binning yielded a pixel scale of  $0''.146\ \text{pixel}^{-1}$ .

The data were reduced and combined using the *gemini* package (V1.10) within IRAF.<sup>8</sup> A master bias frame (constructed by averaging with  $3\sigma$  clipping a series of bias frames) was subtracted from all raw images. Images of the twilight sky were used to flat field the images, then the dithered galaxy images were registered and averaged together with the *imcoadd* task to eliminate the inter-CCD gaps, bad pixels, and cosmic rays.

### 2.2. Radio Observations: Australia Telescope Compact Array

#### 2.2.1. H I

We recreated the 21 cm neutral atomic hydrogen (H I) maps of Dahlem (2005), who kindly provided us with calibrated data, using robust weighting to provide a higher angular resolution

<sup>8</sup> IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

**Table 1**  
Observational Parameters of Radio Observations

	H I	6 cm	3 cm	HCN(1–0)	HCO <sup>+</sup> (1–0)	<sup>13</sup> CO(2–1)	<sup>12</sup> CO(2–1)
Rest frequency (GHz)	1.420	4.786	8.640	88.631	89.188	220.398	230.538
Telescope	ATCA	ATCA	ATCA	ATCA	ATCA	SMA	SMA
Projected baseline (m)	245–6000	153–6000	153–6000	30–200	30–200	16–69	16–69
Field of view	33′	285″	285″	36″	36″	55″	55″
Velocity resolution (km s <sup>−1</sup> )	14	...	...	15	15	18	10
Synthesized beam (arcsec <sup>2</sup> )	20 × 20	2.3 × 1.3	1.1 × 1.1	2.6 × 2.0	2.6 × 2.0	6.9 × 2.8	7.0 × 2.8
P.A. (deg)	0	0	0	80	80	−8.9	−11.9
Reprocessed data (yes/no)	y	y	y	n	n	n	n
Reference	a	b	b	...	...	...	...

**Notes.**

<sup>a</sup> Dahlem (2005).

<sup>b</sup> Forbes et al. (1994a).

than the maps he published. The parameters of the H I observation are summarized in Table 1. The H I maps presented here have a synthesized beam of  $20''.0 \times 20''.0$  at full width at half-maximum (FWHM; by comparison, the H I maps published by Dahlem 2005 have a synthesized beam of  $34''.1 \times 30''.0$ ). In our channel maps, the root-mean-square (rms) noise level in each 13 km s<sup>−1</sup> channel is 1.1 mJy beam<sup>−1</sup>. The integrated H I intensity and intensity-weighted mean H I velocity maps, as well as all other such maps presented in this work, were constructed using the task MOMNT in the AIPS package; this task allows us to select or discard pixels in each channel map based on whether their corresponding intensities when smoothed over a selected number of channels exceed a specified threshold. The H I disk of NGC 7552 is much smaller than the size of the primary beam, and hence no primary beam correction was applied to the data.

### 2.2.2. Radio Continuum

We re-processed the radio continuum data taken by Forbes et al. (1994a) at 3 cm and 6 cm, and re-created their maps to permit our own analyses as well as provide digital versions for use in making overlays. As summarized in Table 1, the observations of Forbes et al. (1994a) were conducted in one 12 hr synthesis on 1993 February 11–12 using the Australia Telescope Compact Array (ATCA) in its 6C configuration. This configuration provides baseline lengths ranging from 153 m to 6000 m, making it possible to image structures (with uniform brightnesses) as large as about  $285''.0$ . The primary flux calibrator used was PKS B1934–638, for which we adopted a flux density of 5.83 Jy at 6 cm (4786 MHz) and 2.84 Jy at 3 cm (8640 MHz). The quasar PKS B2311–452 was used for complex gain calibration.

We calibrated the data in the standard manner using the MIRIAD package. To improve the dynamic range of our maps, we self-calibrated the data sets at both 6 cm and 3 cm from the maps obtained at the respective wavelengths. The final map at 6 cm has a synthesized beam of  $2''.0 \times 1''.3$  at a position angle (P.A.) of 0°, and an rms noise fluctuation of  $72 \mu\text{Jy beam}^{-1}$ . The final map at 3 cm has a synthesized beam of  $1''.1 \times 1''.1$  and an rms noise level of  $66 \mu\text{Jy beam}^{-1}$ . Both the angular resolutions and noise levels of our maps are, of course, similar to those in the maps shown by Forbes et al. (1994a). Given the small angular extent of the circumnuclear radio continuum emission from NGC 7552, no primary beam correction was applied to the data.

### 2.2.3. HCN and HCO<sup>+</sup>

We observed NGC 7552 in HCN (1–0) and HCO<sup>+</sup> (1–0) with the ATCA on three separate occasions between 2005 May and September in the H75, H168, and H214 configurations. As summarized in Table 1, this combination of array configurations provided baseline lengths ranging from 30 m to 200 m, and hence sensitivity to structures (with uniform brightnesses) as large as about  $23''$ . Five of the six antennas of the ATCA were equipped with dual polarization receivers in the 3 mm band, covering a frequency range from 85 GHz to 105 GHz. At the wavelength of the HCN (1–0) (rest frequency of 88.6318470 GHz) and HCO<sup>+</sup> (1–0) (rest frequency of 89.1885230 GHz) lines, the primary beam of the antennas is  $\sim 36''$ . We configured the correlator to observe the HCN (1–0) and HCO<sup>+</sup> (1–0) lines simultaneously in two spectral windows, with the individual windows covered by 64 channels of width 2 MHz (7 km s<sup>−1</sup>) each. Undersampling in the lag domain makes the spectral channels not completely independent; we therefore binned the data into 15 km s<sup>−1</sup> channels for making maps. We adopted a pointing center of  $\alpha_{2000} = 23^{\text{h}}16^{\text{m}}11^{\text{s}}$  and  $\delta_{2000} = -42^{\circ}35'05''$ , coinciding closely with the center of the circumnuclear starburst ring. The typical single-sideband system temperatures varied over the range 250–300 K, and the total on-source integration time was about 17 hr.

We reduced the data in the standard manner using MIRIAD, a procedure that involved baseline corrections, flagging of bad channels and shadowed baselines, calibration, and finally imaging. The absolute flux density scale was set by Uranus or Mars, excluding baselines for which the planets were heavily resolved. We used the bright quasar PKS B1921–293 for bandpass calibration, and the quasar PKS B2326–477 (lying within 5° of our target) for complex gain calibration. We also observed another quasar, PKS B2255–282 (lying within 15° of our target), at frequent intervals to verify the phase solutions.

We made channel maps using robust weighting, yielding a synthesized beam of  $2''.6 \times 2''.0$ , at a P.A. of 80°. The final maps have an rms noise level of 4.5 mJy in HCN (1–0) and 4.2 mJy in HCO<sup>+</sup> (1–0) for each 15 km s<sup>−1</sup> channel. Because the circumnuclear emission has a small angular extent compared with the primary beam, we did not apply any primary beam corrections to the maps.

### 2.3. Radio Observations: Submillimeter Array

We observed NGC 7552 in <sup>12</sup>CO (2–1) and <sup>13</sup>CO (2–1) in the compact configuration of the Submillimeter Array (SMA; Ho et al. 2004) on 2006 August 4. Observations were centered

at  $\alpha_{2000} = 23^{\text{h}}16^{\text{m}}10^{\text{s}}.7$  and  $\delta_{2000} = -42^{\circ}35'05''.41$ , coinciding with the center of the circumnuclear starburst ring as defined by Forbes et al. (1994a). The primary beam of the antennas at the wavelengths of the observed molecular lines is  $55''$ . Seven out of the eight antennas of the SMA were available during our observation. As summarized in Table 1, with baseline lengths ranging from 16 m to 69 m, the largest structure (of uniform brightness) that we expect to be sensitive to is about  $16''$ . The receiver was tuned and the correlator configured to cover the frequency range 227.6–229.6 GHz in the upper sideband (USB) thus including the  $^{12}\text{CO}$  (2–1) line, and 217.6–219.6 GHz in the lower sideband (LSB) thus including the  $^{13}\text{CO}$  (2–1) line. The correlator provided a total of 24 spectral windows (chunks) in each sideband, with a total of 128 channels and a bandwidth of 104 MHz in each chunk. To make the channel maps in  $^{12}\text{CO}$  (2–1), we averaged 10 channels together, resulting in a velocity resolution of  $10 \text{ km s}^{-1}$ . Because the signal is weaker in  $^{13}\text{CO}$  (2–1), to make the channel maps in this line we averaged 18 channels together resulting in a velocity resolution of  $18 \text{ km s}^{-1}$ .

We calibrated the data separately for bandpass, amplitude, and phase using SMA-specific MIR tasks adopted from the MMA software package (written in IDL), which was originally developed for the Owens Valley Radio Observatory (OVRO; Scoville et al. 1993). The bandpass was calibrated using Uranus. For complex gain calibration, we observed the quasars J2235–485 and J2258–279 for a duration of 6 minutes every 30 minutes. We used J2235–485, which is located at an angular distance of  $9:3$  from our target source, for phase calibration. J2235–485, however, is too weak for amplitude calibration. We therefore used J2258–279, which is located at a larger angular distance of  $15^{\circ}$  from our target source, for amplitude calibration. Because the amplitude solutions derived from the individual scans did not differ significantly between adjacent scans, we applied time smoothing scale of three scans of J2258–279 to derive more precise amplitude solutions. We derived the absolute flux calibration scale from Uranus, which we assumed had a flux density of  $36.836 \text{ Jy}$  at the time of our observation.

We imaged the data using MIRIAD. To deconvolve the sidelobes, we first CLEANed the dirty maps without any preselection to judge which features are likely to be real and which are likely to be sidelobes. We then placed boxes around those features judged to be real on the basis that they do not appear to resemble the sidelobes of the dirty beam, and that their structure changed continuously between adjacent channels. As shown in Section 3.5, the features in each channel have a sufficiently simple structure that we are confident of having picked out only those that are real. The  $^{12}\text{CO}$  (2–1) channel maps have an angular resolution of  $7''.0 \times 2''.8$  at P.A. =  $-11:9$ , and a rms noise level of  $78 \text{ mJy beam}^{-1}$  at a velocity resolution of  $10 \text{ km s}^{-1}$ . The  $^{13}\text{CO}$  (2–1) channel maps were made with  $18 \text{ km s}^{-1}$  velocity binning, resulting in an angular resolution of  $6''.9 \times 2''.8$  at a P.A. =  $-8:9$ , and a rms noise level of  $39 \text{ mJy beam}^{-1}$ . Because the features detected span an angular extent from the pointing center that is significantly smaller than the primary beam, we did not apply a primary beam correction to the maps.

### 3. RESULTS

#### 3.1. Optical Images

Figure 1 shows the optical images obtained with Gemini South. Distinct dust lanes are often seen in late-type barred galaxies along the leading edges of the bar, as is the case also

for NGC 7552. In the  $g'$ -band image of Figure 1, we indicate the location of the dust lanes with a more direct tracer, dust emission as imaged by *Spitzer* at  $5.8 \mu\text{m}$ . The overall extent of this dust emission is indicated by a contour, which perfectly matches the dust lanes seen in the optical.

The  $\text{H}\alpha$  image of Figure 1 highlights regions of recent activity involving massive stars. The circumnuclear starburst ring as traced in the 3 cm continuum is overlaid on the  $\text{H}\alpha$  image to show the overall extent of this ring. The physical size of the ring at 3 cm is comparable to the size of the central bright bulge of NGC 7552 in the optical images. A magnified view of the central region in  $\text{H}\alpha$  and 3 cm continuum is shown in Figure 3. In addition to the circumnuclear starburst ring, a bright  $\text{H II}$  region, detected in both  $\text{H}\alpha$  and radio continuum, is apparent  $\sim 3.5 \text{ kpc}$  east of the nucleus along the bar. This  $\text{H II}$  region was also detected in earlier  $\text{H}\alpha$  images by Forbes et al. (1994a).

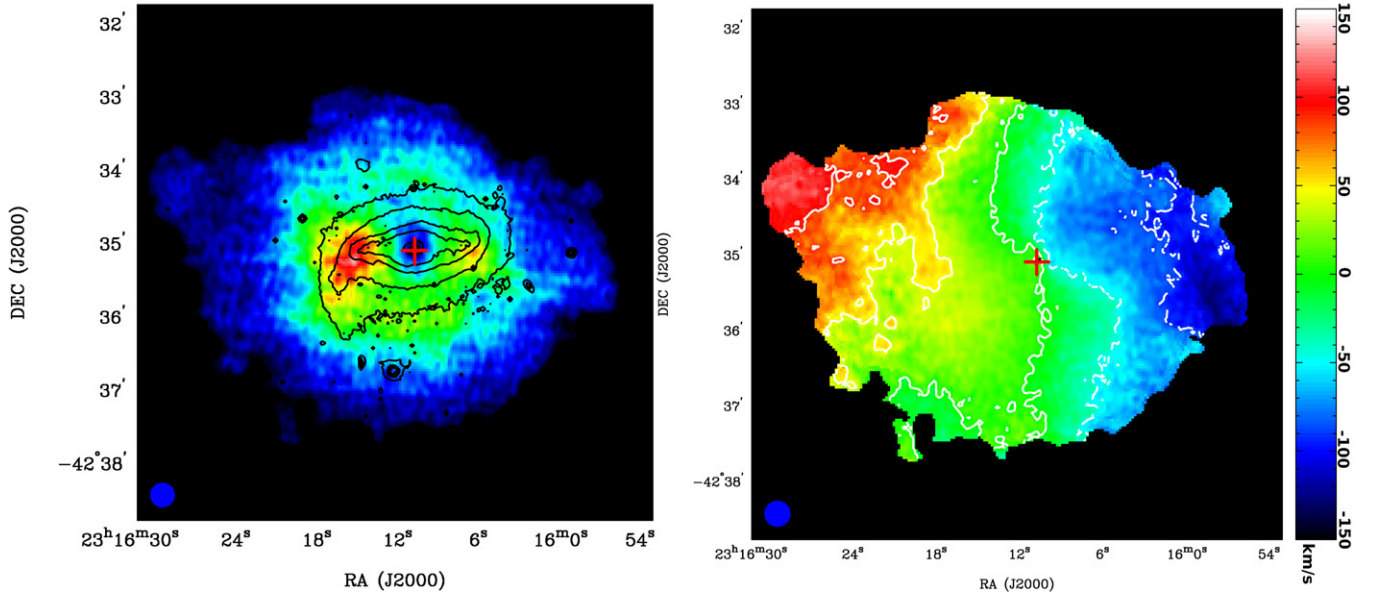
#### 3.2. $\text{H I}$

A map of the integrated  $\text{H I}$  intensity (zeroth moment) is shown in Figure 2 (left panel). As pointed out by Dahlem (2005), the  $\text{H I}$  gas traces the large-scale disk of NGC 7552. For ease of comparison, the optical  $i'$ -band image of the galaxy shown earlier in Figure 1 is overlaid as contours. Along the major axis the  $\text{H I}$  disk can be detected out to a projected radius of about  $2100''$  ( $\sim 20 \text{ kpc}$ ), about twice as far out as the optical disk. Similarly, along the minor axis, the  $\text{H I}$  disk can be detected out to a projected radius of about  $121''$  ( $\sim 13 \text{ kpc}$ ), again about twice as far out as the optical disk. There is a prominent ring-like enhancement in  $\text{H I}$  that peaks at or close to the ends of the large-scale optical bar. Finally, a central depression can be seen in the integrated  $\text{H I}$  intensity map. This central depression has a size that closely coincides with the angular extent of the circumnuclear molecular gas detected in our observations as described below (Sections 3.4 and 3.5).

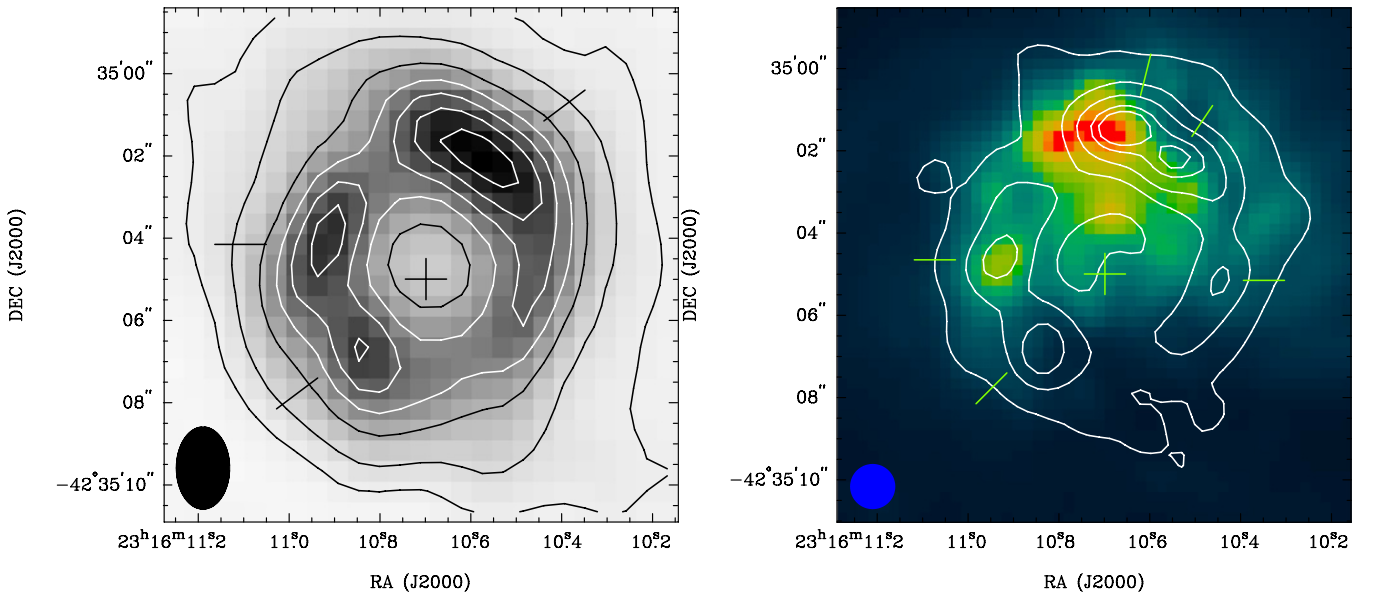
Figure 2 (right panel) shows a map of the intensity-weighted mean  $\text{H I}$  velocity (first moment). Assuming a systemic velocity of  $1586 \text{ km s}^{-1}$  as found in the optical and as derived below (Section 3.5) from our molecular line observations, emission can be seen spanning the velocity range from  $-100 \text{ km s}^{-1}$  to  $100 \text{ km s}^{-1}$  about the systemic velocity. As can be seen, the kinematic major axis is aligned roughly east–west, coincident with the major axis of the optical bar. Blueshifted emission (negative velocities) originates primarily westward and redshifted emission (positive velocities) primarily eastward of the galactic center. On close scrutiny, the P.A. of the kinematic major axis can be seen to change with galactic radius, creating an S-shaped pattern in the gas kinematics. Such S-shaped kinematic patterns are commonly attributed to non-circular motions under the non-axisymmetric gravitational potential of a galactic bar (see Section 5.1 for detailed discussion). Far beyond the optical bar however, the P.A. of the kinematic major axis continues to change, suggesting that the outskirts of the galaxy may be warped (by gravitational interactions with one or more neighboring galaxies as described in Section 1).

#### 3.3. Radio Continuum

Figure 3 shows the central part of our radio continuum images at 6 cm (left panel) and 3 cm (right panel), both of which were recreated from the data taken by Forbes et al. (1994a). Over the field shown, the radio continuum emission originates from a ring, which as mentioned by Forbes et al. (1994a) corresponds to the circumnuclear starburst ring, as can be seen in Figure 3



**Figure 2.** Integrated H I intensity (left panel) and intensity-weighted mean H I velocity (right panel), both in color. In the left panel, contours are of the  $i'$  image shown in Figure 1. In the right panel, contours correspond to the intensity-weighted mean H I velocity, just like the color map, and are plotted in steps of  $-90$ ,  $-45$ ,  $0$ ,  $45$ , and  $90$   $\text{km s}^{-1}$  with respect to the systemic velocity. In each panel, a cross is plotted at the center of the galaxy, corresponding to the center of the starburst ring (Forbes et al. 1994a), and the synthesized beam shown at the lower left corner.

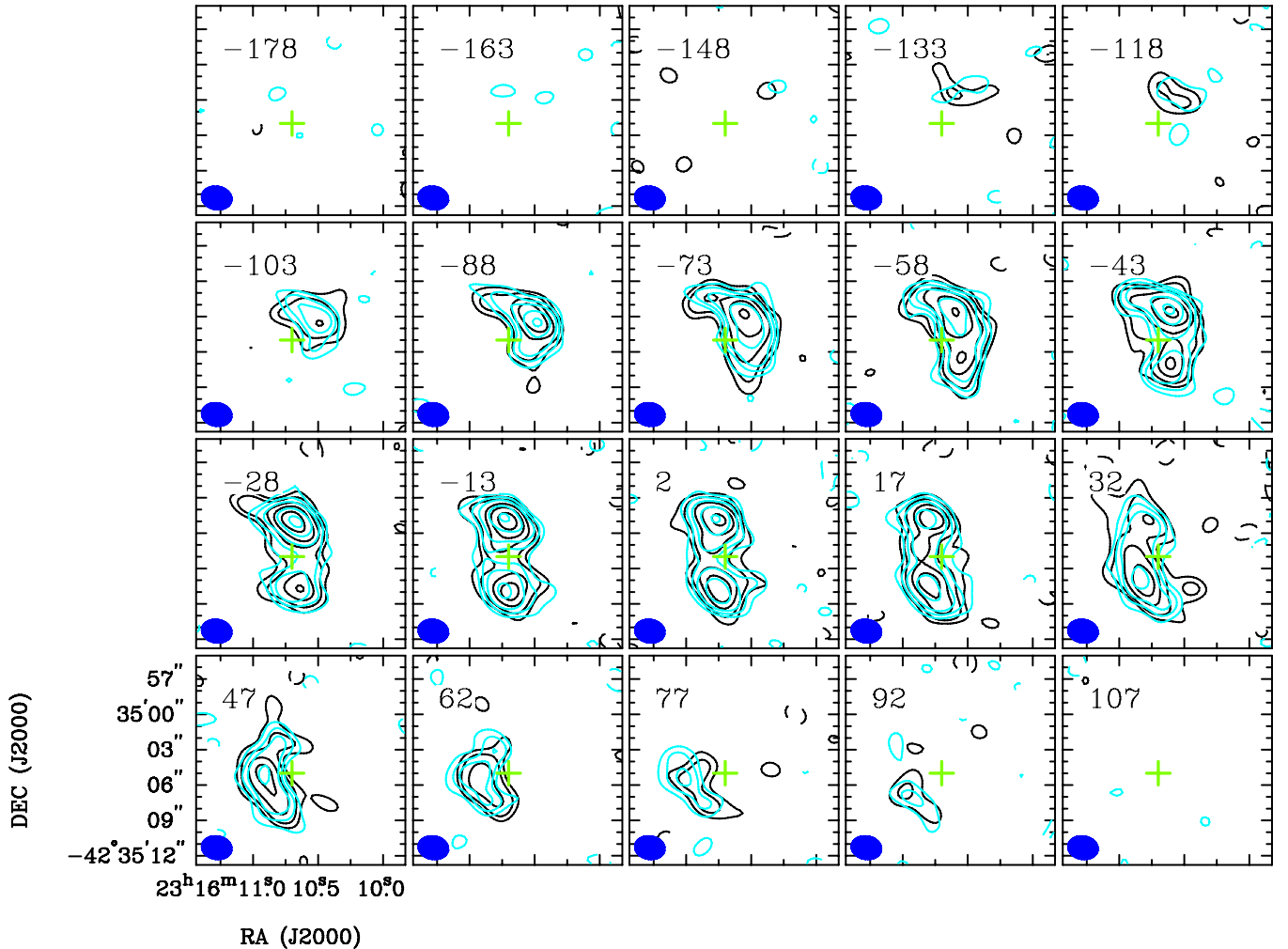


**Figure 3.** Radio continuum images of the circumnuclear starburst ring in NGC 7552 at 6 cm (left panel; both contours and gray scale) and 3 cm (right panel; contours recreated from the data taken by and published in Forbes et al. (1994a)). In the right panel, the 3 cm (contours) is overlaid on the  $H\alpha$  (color scale) image shown in Figure 1. Straight lines in both panels indicate peaks in radio emission, which correspond to star-forming knots. Contour levels of the 6 cm image are at 3, 10, 25, 40, 55, 70, and  $85\sigma$ , where  $1\sigma = 72 \mu\text{Jy beam}^{-1}$ , and contour levels of the 3 cm image are at 3, 5, 15, 25, 35, and  $40\sigma$ , where  $1\sigma = 66 \mu\text{Jy beam}^{-1}$ . A cross in each panel indicates the center of the ring as defined by Forbes et al. (1994a). The synthesized beam is  $2''.0 \times 1''.3$  at 6 cm and  $1''.1 \times 1''.1$  at 3 cm, and shown at the lower left corner of the respective panels.

(right panel) where the 3 cm image (contours) is overlaid on our  $H\alpha$  image (color). This ring has an outer extent comparable to the central H I depression mentioned above.

Five peaks are clearly visible in the 3 cm image (right panel), whereas only three peaks are visible in the 6 cm image (left panel) that has a lower angular resolution than the 3 cm image. To study the spectral index of each knot, we have convolved the 3 cm image to the same angular resolution as the 6 cm image. After degrading the angular resolution, we find that only three

knots are discernible in the 3 cm image, which now resembles the 6 cm image. The spectral indices ( $\alpha$ , where  $S \propto \nu^\alpha$  with  $S$  the flux density and  $\nu$  the observing frequency) of the three knots (see the left panel of Figure 3) are  $-1.11 \pm 0.02$  (upper right),  $-0.84 \pm 0.03$  (upper left), and  $-0.72 \pm 0.05$  (bottom left), with an average value of  $-0.89 \pm 0.06$  that is similar to that inferred by Forbes et al. (1994a). The negative spectral indices indicate that the emission of all the knots is dominated by, if not produced entirely by (non-thermal), synchrotron emission.



**Figure 4.** Channel maps of the central region of NGC 7552 in HCN (1–0) (black contours) and HCO<sup>+</sup> (1–0) (cyan contours). Contour levels are plotted at  $-2, 2, 3, 5, 7,$  and  $10 \times 4.2 \text{ mJy beam}^{-1}$  (rms noise level) for HCN (1–0) and  $\times 4.5 \text{ mJy beam}^{-1}$  (rms noise level) for HCO<sup>+</sup> (1–0). The velocity of each channel, in units of  $\text{km s}^{-1}$ , is shown in the upper left corner of each panel. A cross is plotted at the center of the galaxy in each panel. The synthesized beam is  $2''.6 \times 2''.0$  at a P.A. of  $80^\circ$ , and is shown at the lower left corner of each panel.

(A color version of this figure is available in the online journal.)

### 3.4. HCN and HCO<sup>+</sup>

Figure 4 shows our channel maps of the HCN (black contours) and HCO<sup>+</sup> (1–0) (cyan contours) lines. Both HCN (1–0) and HCO<sup>+</sup> (1–0) emission are concentrated in two knots located to the north and south of the galactic nucleus. The northern knot is detectable over a velocity range from  $-133 \text{ km s}^{-1}$  to  $32 \text{ km s}^{-1}$  with respect to the optically determined systemic velocity of the galaxy, and is therefore preferentially blueshifted. The southern knot is detectable over a velocity range from about  $-43 \text{ km s}^{-1}$  to  $92 \text{ km s}^{-1}$  with respect to the systemic velocity, and is therefore preferentially redshifted. The centroids of both knots shift from west to east with increasing (i.e., from blueshifted to redshifted) velocities.

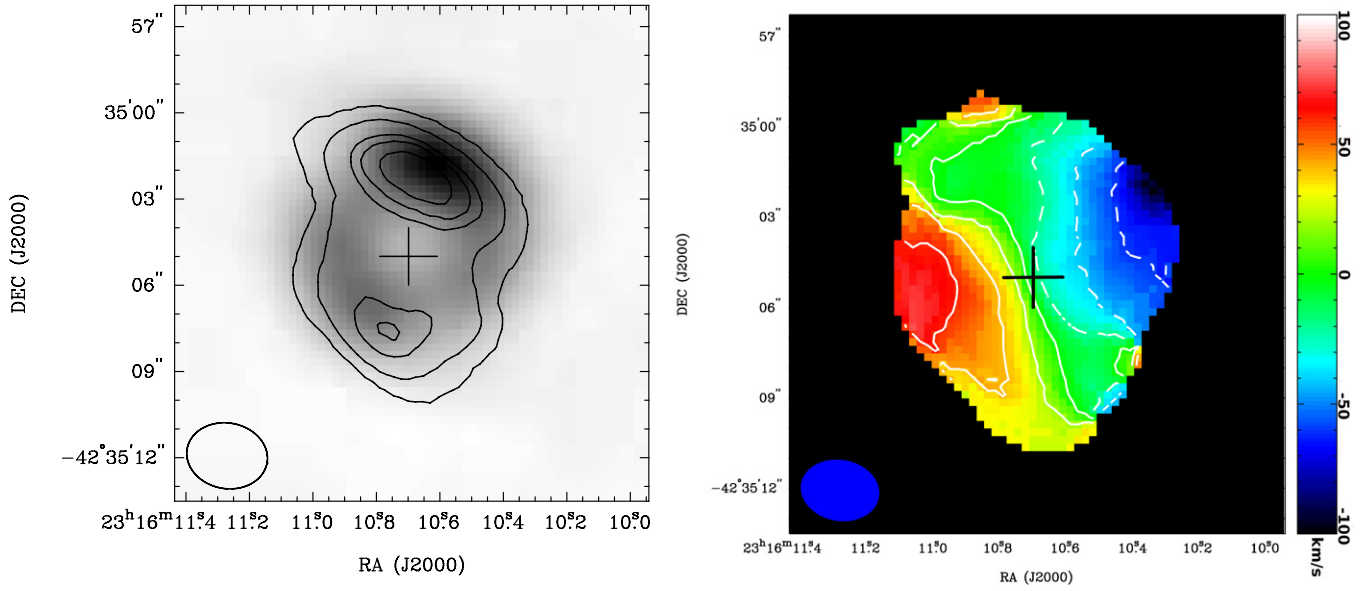
Figure 5 (left panel) shows a map of the integrated HCN (1–0) intensity in black contours. The overall size of the emitting region is comparable with the central H I depression described above in Section 3.2, suggesting that the gas density is so high here that the atomic hydrogen gas has been largely converted to molecular hydrogen gas. The HCN (1–0) contours are overlaid on a gray-scale image of the 3 cm continuum emission shown earlier in Figure 3 (right panel) but now convolved to the same

angular resolution as the HCN (1–0) map. At this angular resolution, the 3 cm continuum map shows two knots just like the HCN (1–0) map. The observed HCN (1–0) emission therefore originates from the circumnuclear starburst ring, and exhibits a central depression or possibly a hole. Notice the small displacement between the centroids of the HCN (1–0) gas and radio continuum knots, a point we shall return to in Section 6.2.

Figure 5 (right panel) shows a map of the intensity-weighted mean HCN (1–0) velocity. A velocity gradient is discernible roughly along the east–west direction, with the kinematic major axis lying at a P.A. of  $110^\circ$ . The kinematic major axis in HCN (1–0), as well as in HCO<sup>+</sup> (1–0) as described below, are both aligned with the kinematic major axis at the innermost region of the large-scale H I disk as described in Section 3.2. The good agreement suggests that there is a continuous velocity field between the outermost region in molecular gas and innermost region in neutral atomic gas.

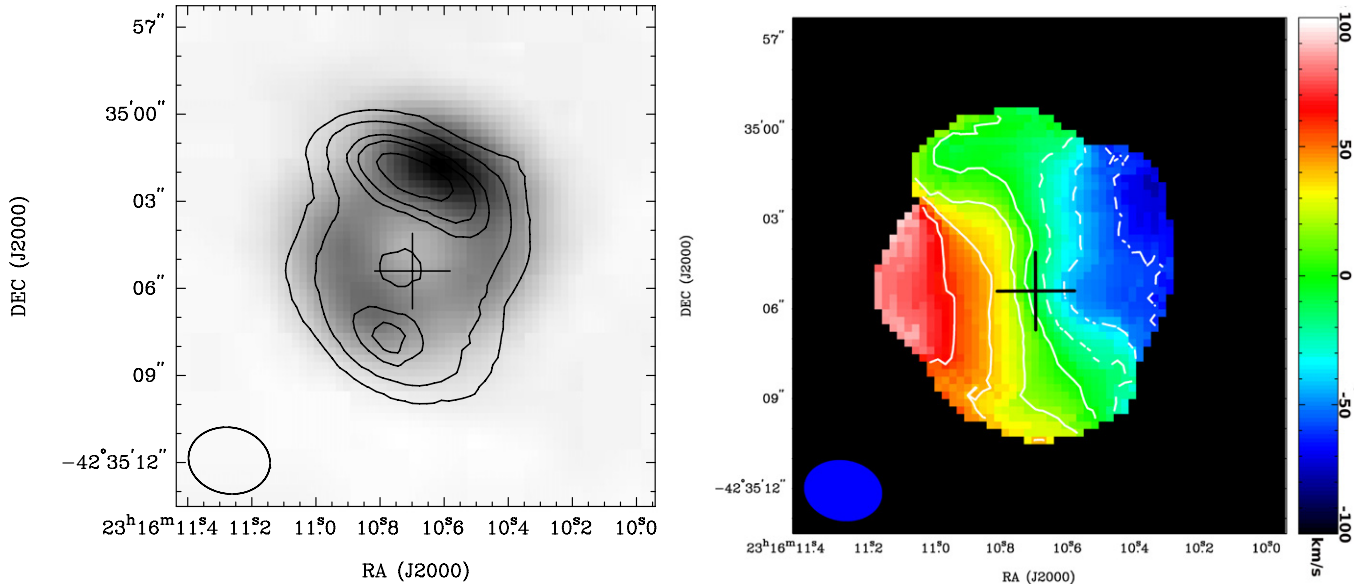
Figure 6 shows our moment maps of the HCO<sup>+</sup> (1–0) line. The emission in all these maps is virtually identical to that seen in HCN (1–0). Both HCN (1–0) and HCO<sup>+</sup> (1–0) have comparable critical densities of  $\sim 10^6 \text{ cm}^{-3}$ , suggesting that both are tracing the same regions of dense gas.





**Figure 5.** Integrated HCN (1–0) intensity (left panel; contours) and intensity-weighted mean HCN (1–0) velocity (right panel; contours and color scale) derived from the channel maps shown in Figure 4. In the left panel, contours indicating the integrated HCN (1–0) intensity are plotted at 3, 6, 10, 12, and  $15\sigma$ , where  $1\sigma = 0.3 \text{ Jy beam}^{-1} \text{ km s}^{-1}$ . Gray scale corresponds to the 3 cm continuum image shown in Figure 3, but convolved to the same angular resolution as the HCN (1–0) image. In the right panel, contours correspond to the intensity-weighted mean HCN (1–0) velocity, just like the color map, and are plotted in steps of  $20 \text{ km s}^{-1}$  starting from  $-60 \text{ km s}^{-1}$  and ending at  $60 \text{ km s}^{-1}$  with respect to the systemic velocity. A cross is plotted at the center of the galaxy, and the synthesized beam is shown at the lower left corner, of each panel.

(A color version of this figure is available in the online journal.)



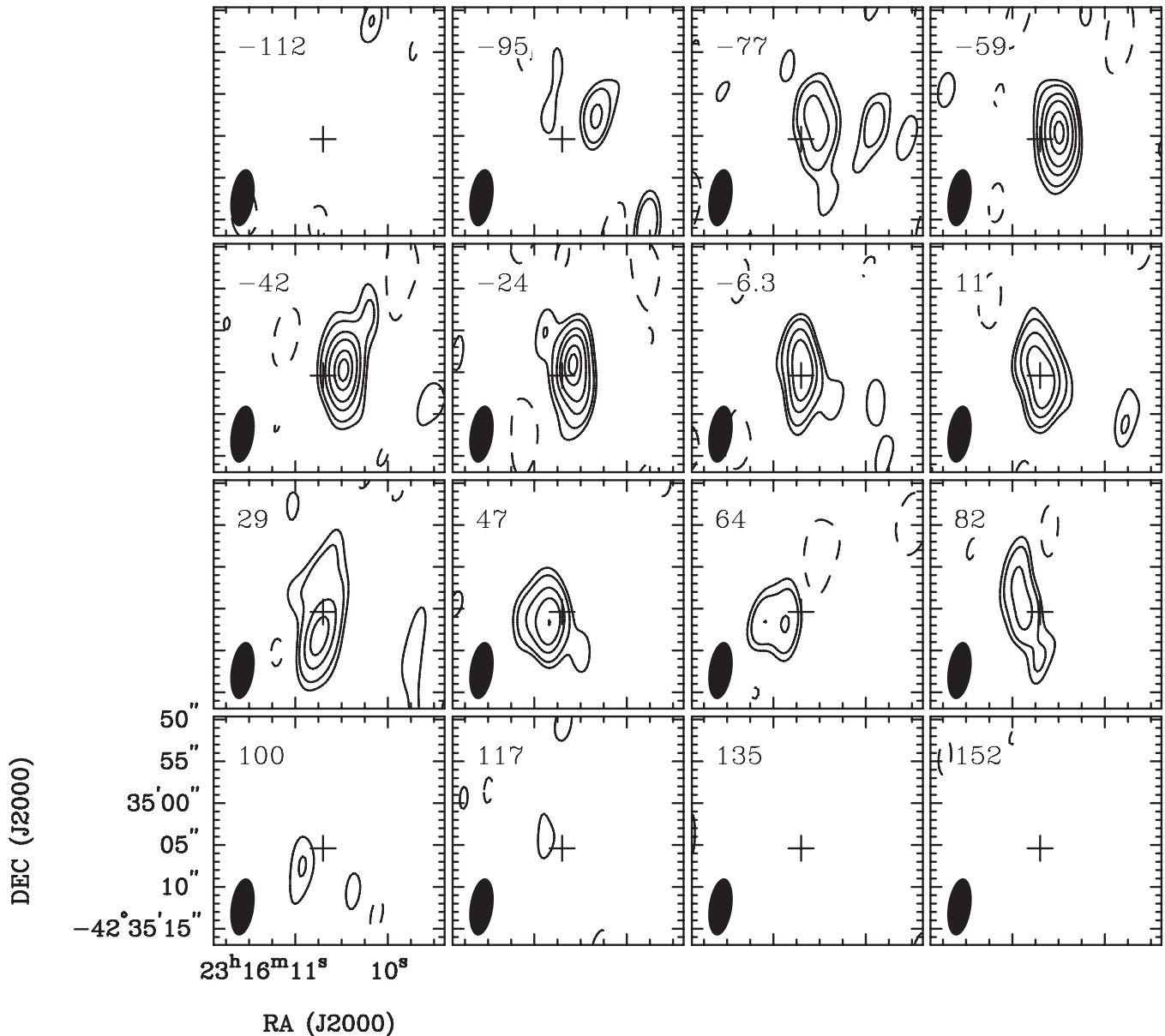
**Figure 6.** Integrated  $\text{HCO}^+$  (1–0) intensity (left panel; contours) and intensity-weighted mean  $\text{HCO}^+$  (1–0) velocity (right panel; contours and color scale) derived from the channel maps shown in Figure 4. In the left panel, contours indicating the integrated  $\text{HCO}^+$  (1–0) intensity are plotted at 3, 6, 10, 12, and  $15\sigma$ , where  $1\sigma = 0.3 \text{ Jy beam}^{-1} \text{ km s}^{-1}$ . Gray scale corresponds to the 3 cm continuum image shown in Figure 3, but convolved to the same angular resolution as the  $\text{HCO}^+$  (1–0) image. In the right panel, contours correspond to the intensity-weighted mean  $\text{HCO}^+$  (1–0) velocity, just like the color map, and are plotted in steps of  $20 \text{ km s}^{-1}$  starting from  $-60 \text{ km s}^{-1}$  and ending at  $60 \text{ km s}^{-1}$  with respect to the systemic velocity. A cross is plotted at the center of the galaxy, and the synthesized beam is shown at the lower left corner, of each panel.

(A color version of this figure is available in the online journal.)

### 3.5. CO

Figure 7 shows our channel maps of the  $^{13}\text{CO}$  (2–1) line, which by contrast with both the HCN (1–0) and  $\text{HCO}^+$  (1–0) has a much lower critical density of just  $\sim 10^3 \text{ cm}^{-3}$ . Only a single compact component is discernible that shifts from west to east across the center of the galaxy with increasing velocity, just like that seen in HCN (1–0) and  $\text{HCO}^+$  (1–0) as described

above in Section 3.4. We have convolved both the HCN(1–0) and  $\text{HCO}^+$  (1–0) maps to the same angular resolution as the  $^{13}\text{CO}$  (2–1) map, and found that all show the same spatial-kinematic structure suggesting that they all trace the same global features. In the  $^{13}\text{CO}$  (2–1) map, emission is detectable over the velocity range from about  $-95 \text{ km s}^{-1}$  to  $100 \text{ km s}^{-1}$  with respect to the systemic velocity, comparable to the detectable velocity range of the emission in both the HCN (1–0) and  $\text{HCO}^+$  (1–0) maps.



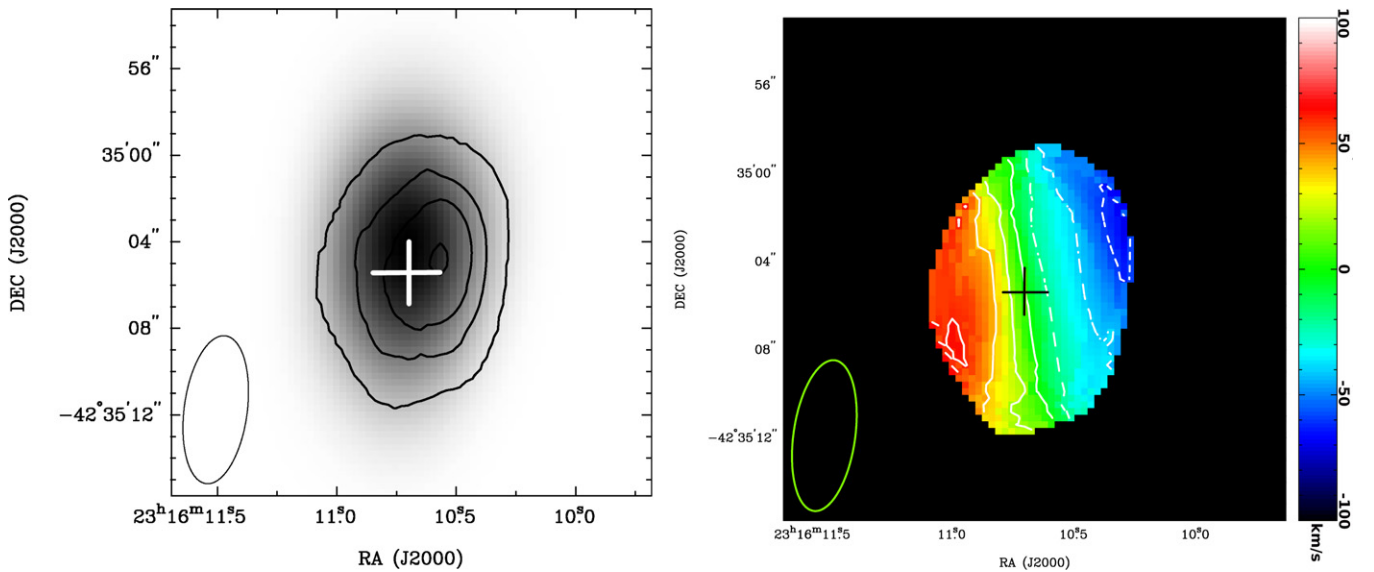
**Figure 7.** Channel maps of the central region of NGC 7552 in  $^{13}\text{CO}$  (2–1). Contour levels are plotted at  $-2, 2, 3, 5, 7, 10,$  and  $12 \times 39 \text{ mJy beam}^{-1}$  (rms noise level). The velocity of each channel, in units of  $\text{km s}^{-1}$ , is shown in the upper left corner of each panel. A cross is plotted at the center of the galaxy in each panel. The synthesized beam is  $6.9 \times 2.8$  at a P.A. of  $-8.1$ , and is shown at the lower left corner of each panel.

Figure 8 (left panel) shows the integrated  $^{13}\text{CO}$  (2–1) intensity map in black contours overlaid on the integrated HCN (1–0) intensity map, taken from Figure 5 but convolved to the same angular resolution as the  $^{13}\text{CO}$  (2–1) map, in gray scale. Figure 8 (right panel) shows a map of the intensity-weighted mean  $^{13}\text{CO}$  (2–1) velocity. As can be seen, the compact central component detected in  $^{13}\text{CO}$  (2–1) has a similar size and kinematics as the circumnuclear HCN (1–0) (and  $\text{HCO}^+$  (1–0)) emission. In Figure 9, we show position–velocity (P–V) diagrams of the HCN (1–0),  $\text{HCO}^+$  (1–0), and  $^{13}\text{CO}$  (2–1) emission along a P.A. of  $110^\circ$ , which corresponds to the orientation of the kinematic major axis. As can be seen, the P–V diagram in  $^{13}\text{CO}$  (2–1) closely resembles that in the higher density tracers of molecular hydrogen gas. We therefore believe that the observed  $^{13}\text{CO}$  (2–1) emission originates from essentially the same regions of the circumnuclear starburst ring as the observed HCN (1–0) or  $\text{HCO}^+$  (1–0) emission. Of course, any central depression or

hole of the same size as that seen in HCN (1–0) and  $\text{HCO}^+$  (1–0) would not be discernible at the angular resolution of our  $^{13}\text{CO}$  (2–1) observation.

The P–V diagrams of Figure 9 reveal that the systemic velocity at the center of the circumnuclear starburst ring, which presumably coincides with the dynamical center of the galaxy, is about  $1580 \text{ km s}^{-1}$ . This is consistent with the optically determined value of  $1586 \pm 5 \text{ km s}^{-1}$  (RC3), but significantly smaller than that determined from the global H I profile of  $1608 \pm 5 \text{ km s}^{-1}$  (The H I Parkes All Sky Survey; Barnes et al. 2001).

Figure 10 shows our channel maps of the  $^{12}\text{CO}$  (2–1) line. By contrast with the simplicity of the  $^{13}\text{CO}$  (2–1) channel maps, three components are now discernible: a compact central component that is the counterpart of that seen in  $^{13}\text{CO}$  (2–1), and two extensions on opposite sides of the center with one arm extending east from north and the other arm west from



**Figure 8.** Integrated  $^{13}\text{CO}$  (2–1) intensity (left panel; contours) and intensity-weighted mean  $^{13}\text{CO}$  (2–1) velocity (right panel; contours and color scale) derived from the channel maps shown in Figure 7. In the left panel, contours indicating the integrated  $^{13}\text{CO}$  (2–1) intensity are plotted at  $3\sigma$ ,  $10\sigma$ ,  $14\sigma$ , and  $18\sigma$ , where  $1\sigma = 2.3 \text{ Jy beam}^{-1} \text{ km s}^{-1}$ . Gray scale corresponds to the integrated HCN (1–0) intensity image shown in Figure 5, but convolved to the same angular resolution as the  $^{13}\text{CO}$  (2–1) image. In the right panel, contours correspond to the intensity-weighted mean  $^{13}\text{CO}$  (2–1) velocity, just like the color map, and are plotted in steps of  $20 \text{ km s}^{-1}$  starting from  $-60 \text{ km s}^{-1}$  and ending at  $60 \text{ km s}^{-1}$  with respect to the systemic velocity. A cross is plotted at the center of the galaxy, and the synthesized beam is shown at the lower left corner, of each panel.

(A color version of this figure is available in the online journal.)

south. The compact central component is detectable over the velocity range from about  $-125 \text{ km s}^{-1}$  to  $135 \text{ km s}^{-1}$  with respect to the systemic velocity, somewhat larger than the detectable velocity range of its counterparts in  $^{13}\text{CO}$  (2–1), HCN (1–0), and  $\text{HCO}^+$  (1–0). By contrast, both the eastern and western arm-like components are detectable over a narrower velocity range than the compact central component in any of the observed molecular lines. The western arm is detectable over the velocity range from about  $-85 \text{ km s}^{-1}$  to  $55 \text{ km s}^{-1}$  and is therefore preferentially blueshifted, whereas the eastern arm is detectable over the velocity range from about  $-55 \text{ km s}^{-1}$  to  $75 \text{ km s}^{-1}$  and is therefore preferentially redshifted.

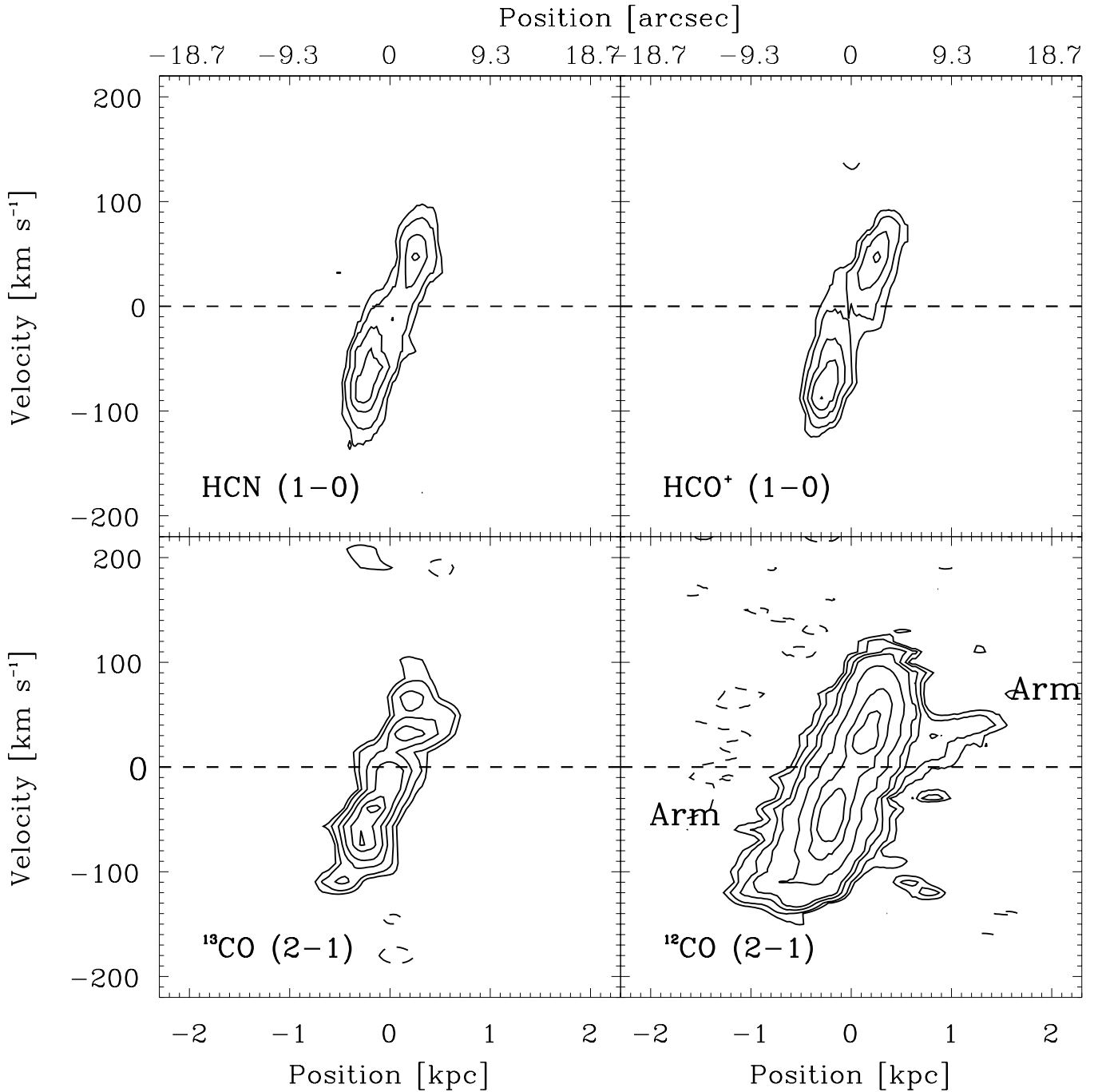
Figure 11 (upper left panel) shows a map of the integrated  $^{12}\text{CO}$  (2–1) intensity in black/white contours overlaid on the optical  $i'$ -band image of the inner part of the galaxy. Dust lanes are visible as silhouettes in the optical image running east and west of the center, as well as in emission at  $5.8 \mu\text{m}$  as shown by the cyan contour based on the observations by Kennicutt et al. (2003) in their *Spitzer* SINGS project. The eastern and western  $^{12}\text{CO}$  (2–1) arms coincide with the innermost regions of these dust lanes. These arms appear to connect with the circumnuclear starburst ring at or near the locations where we detected two knots of dense molecular gas as traced in HCN (1–0) and  $\text{HCO}^+$  (1–0) (see Section 3.4), as shown in the lower left panel of Figure 11 where the integrated  $^{12}\text{CO}$  (2–1) intensity is plotted in green contours and integrated HCN (1–0) intensity in black contours. We note that, for the same ratio in intensity between the arms and compact central component, the stronger western arm should be detectable (at  $4.5\sigma$ ) at the sensitivity of our HCN (1–0)/ $\text{HCO}^+$  (1–0) maps.

Figure 11 (lower right panel) shows a map of the intensity-weighted mean  $^{12}\text{CO}$  (2–1) velocity. At our angular resolution, the kinematics of the compact central component cannot be completely separated from the individual kinematics of the

eastern and western arms, and vice versa. Nevertheless, the S-shape velocity field (contours) indicates the existence of non-circular motion, which is similar to what is observed in other barred galaxies. In Figure 9, we show a P–V diagram made from the CO (2–1) channel maps at a P.A. of  $110^\circ$ . As can be seen, the kinematics of the compact central component are similar to its counterparts in the other observed molecular lines. The western arm is preferentially blueshifted whereas the eastern arm preferentially redshifted, and therefore both arms share the same overall kinematics as the large-scale galactic disk seen in HI.

We derive the mass of molecular hydrogen gas in all the three features detected in  $^{12}\text{CO}$  (2–1) in the standard manner using the Galactic conversion factor between CO intensity and mass of molecular hydrogen gas of  $X_{\text{CO}} = 3 \times 10^{20} \text{ cm}^{-2} (\text{K km s}^{-1})^{-1}$  (Solomon et al. 1987). This conversion is based on the  $^{12}\text{CO}$  (1–0) line, and in the present situation requires knowledge of the ratio in  $^{12}\text{CO}$  (2–1) to  $^{12}\text{CO}$  (1–0) intensities. Aalto et al. (1995) find a ratio in brightness temperatures between these two lines of  $1.2 \pm 0.1$  based on single-dish observations with the Swedish-ESO 15 m Submillimeter Telescope (SEST) toward the center of NGC 7552. For simplicity, we shall therefore assume a ratio in brightness temperature of unity for  $^{12}\text{CO}$  (2–1) to  $^{12}\text{CO}$  (1–0) throughout this work. Accordingly, the compact central component (corresponding to the circumnuclear starburst ring, and perhaps regions within), which has an integrated intensity in  $^{12}\text{CO}$  (2–1) of  $1487.1 \pm 20.3 \text{ Jy km s}^{-1}$ , has a gas mass of  $(3.2 \pm 0.1) \times 10^9 M_\odot$ , including the mass of the heavy elements other than molecular hydrogen ( $M_{\text{gas}} = 1.36 M_{\text{H}_2}$ ). The integrated  $^{12}\text{CO}$  (2–1) intensity detected over the entire region is  $2330.5 \pm 34.7 \text{ Jy km s}^{-1}$ , and hence the mass of molecular gas in the eastern and western arms combined is  $(1.7 \pm 0.2) \times 10^9 M_\odot$ .

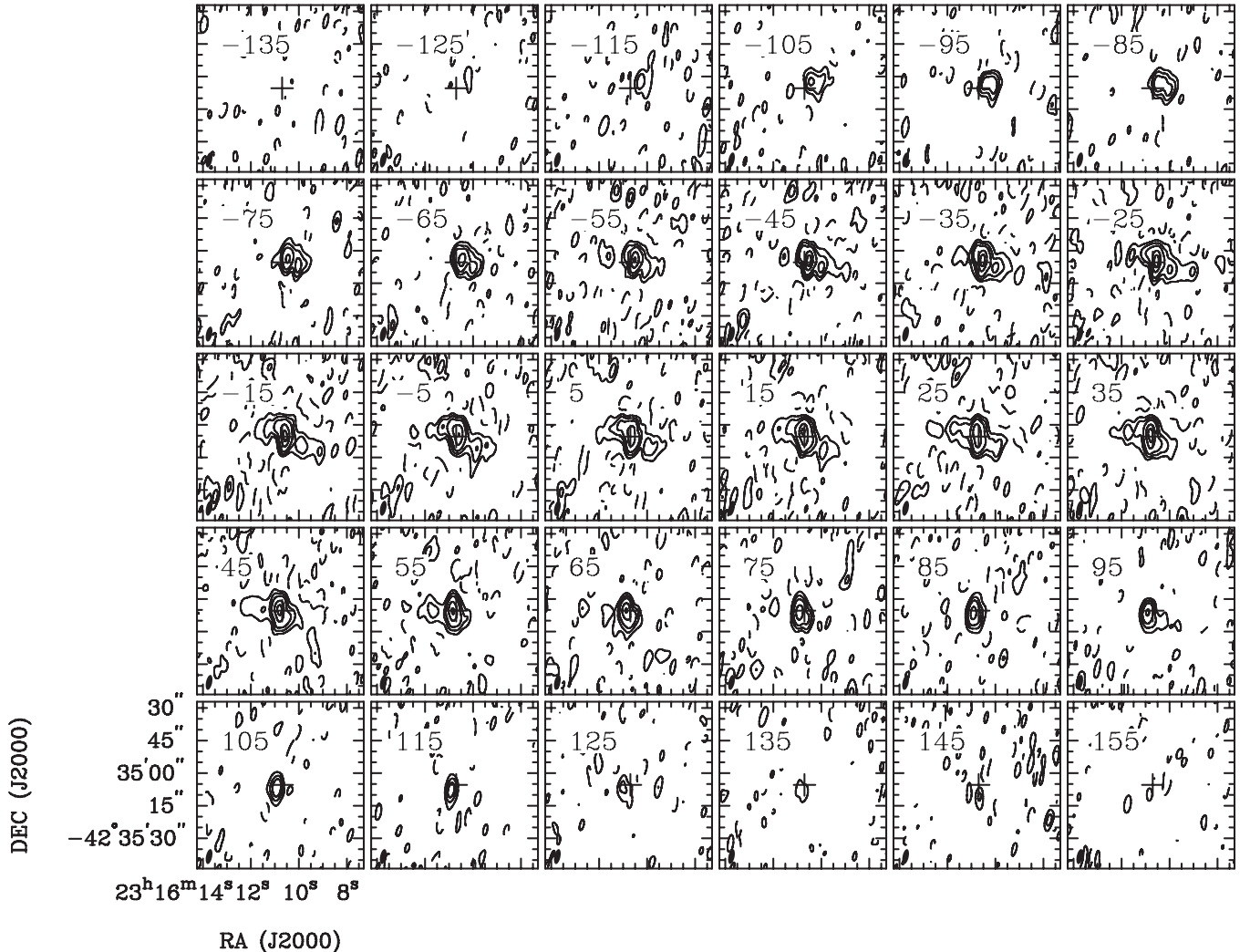
The dense gas fraction can be estimated by comparing the mass derived from  $^{12}\text{CO}$  (2–1) and HCN (1–0).



**Figure 9.** Position–velocity (P–V) diagrams of HCN (1–0) (upper left panel), HCO<sup>+</sup> (1–0) (upper right panel), <sup>13</sup>CO (2–1) (lower left panel), and <sup>12</sup>CO (2–1) (lower right panel) along a P.A. of 110°, which corresponds to the orientation of the kinematic major axis. Negative positions are located eastward and positive positions located westward of the galaxy center. Dashed horizontal lines indicate 0 km s<sup>−1</sup> relative to the systemic velocity of 1580 km s<sup>−1</sup>. Contour levels of HCN (1–0) and HCO<sup>+</sup> (1–0) are at −2, 2, 3, 5, 7, and 10 $\sigma$ , where 1 $\sigma$  is 4.2 and 4.5 mJy beam<sup>−1</sup>, respectively, for HCN (1–0) and HCO<sup>+</sup> (1–0). Contours levels of <sup>13</sup>CO (2–1) are at −2 $\sigma$ , 2 $\sigma$ , 3 $\sigma$ , 5 $\sigma$ , 7 $\sigma$ , 10 $\sigma$ , 12 $\sigma$ , and 15 $\sigma$ , where 1 $\sigma$  = 39 mJy beam<sup>−1</sup>, and <sup>12</sup>CO (2–1) at −2 $\sigma$ , 2 $\sigma$ , 3 $\sigma$ , 5 $\sigma$ , 7 $\sigma$ , 15 $\sigma$ , 30 $\sigma$ , and 40 $\sigma$ , where 1 $\sigma$  = 78 mJy beam<sup>−1</sup>.

Assuming that the gas traced by HCN (1–0) is gravitationally bound (Solomon et al. 1992), the HCN-to-H<sub>2</sub> conversion factor  $X_{\text{HCN}} \equiv N(\text{H}_2)/L'_{\text{HCN}}$  is  $\sim 20 M_{\odot} (\text{K km s}^{-1} \text{pc}^{-2})^{-1}$ , where  $N(\text{H}_2)$  is the column density of molecular hydrogen gas and  $L'_{\text{HCN}}$  the HCN (1–0) luminosity. For the measured  $L'_{\text{HCN}} \approx 7.5 \times 10^7 \text{ K km s}^{-1} \text{pc}^{-2}$ , the corresponding mass in dense gas, as traced in HCN (1–0), is  $\sim 1.5 \times 10^9 M_{\odot}$ . The dense gas fraction is therefore about 50% at the compact central component, much higher than is usually found in Galactic GMCs. We note, however, that the inferred dense gas fraction

may contain considerable uncertainty because the conversion factors  $X_{\text{HCN}}$  and  $X_{\text{CO}}$  used may not be appropriate for the circumnuclear regions of galaxies (although the use of different conversion factors can either raise or lower the dense gas fraction). For example, based on virial arguments for the circumnuclear starburst ring of the nearby barred spiral galaxy IC 342, Meier & Turner (2001) argue that  $X_{\text{CO}}$  is only  $1 \times 10^{20} \text{ cm}^{-2} (\text{K km s}^{-1})^{-1}$  or even smaller, in which case the dense gas fraction at the compact central component becomes nearly 100%.



**Figure 10.** Channel maps of the central region of NGC 7552 in  $^{12}\text{CO}$  (2–1). Contour levels are plotted at  $-2, 2, 5, 10, 25, 40,$  and  $55 \times 78 \text{ mJy beam}^{-1}$  (rms noise level). The velocity of each channel, in units of  $\text{km s}^{-1}$ , is shown in the upper left corner of each panel. A cross is plotted at the center of the galaxy in each panel. The synthesized beam is  $7''.0 \times 2''.8$  at a P.A. of  $-11^\circ 9'$ , and is shown at the lower left corner of each panel.

#### 4. PHYSICAL PROPERTIES OF THE CIRCUMNUCLEAR MOLECULAR GAS

In this section, we derive the physical properties of the molecular gas in the circumnuclear starburst ring based on the ratio in brightness temperatures of the molecular lines observed. We then investigate the dynamical stability of this gas.

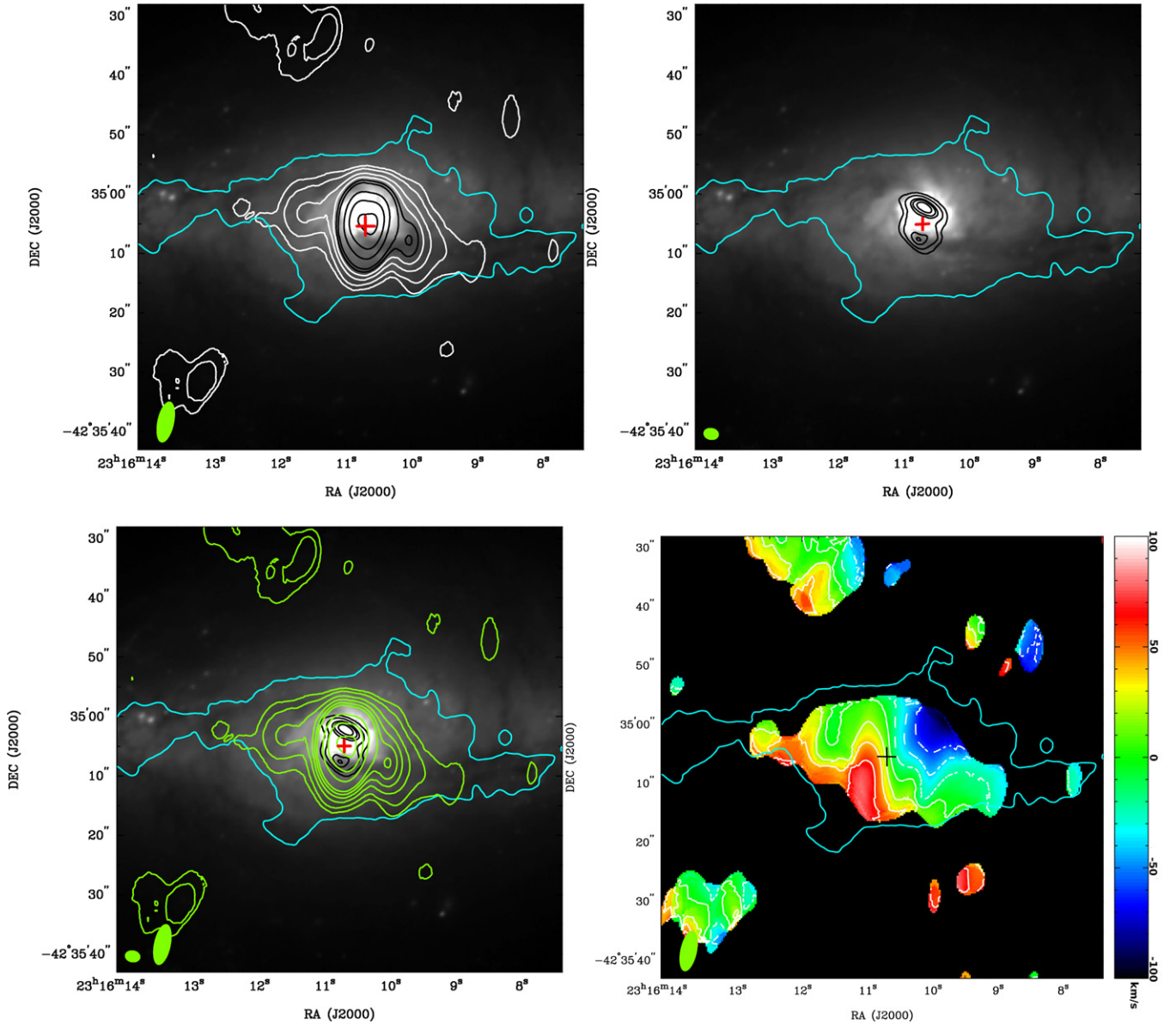
##### 4.1. Emission Recovered in Our Interferometric Observations

Line ratios inferred with an interferometer can be compromised if there exist uniformly bright features with sizes beyond that detectable even by the shortest baseline. Aalto et al. (1995) have observed the central region of NGC 7552 with the SEST in  $^{12}\text{CO}$  (1–0),  $^{12}\text{CO}$  (2–1),  $^{13}\text{CO}$  (1–0),  $^{13}\text{CO}$  (2–1), and HCN (1–0). They do not show line profiles for any of the transitions observed, but only quote the line intensities or ratio in brightness temperatures of the different lines. After convolving our channel maps to the same angular resolution as in the observation of the corresponding line with the SEST, we found that our SMA observation recovered  $54\% \pm 9\%$  of the  $^{12}\text{CO}$  (2–1) emission detected by the SEST. Given typical systematic uncertainties of about 20% in the absolute flux calibration at millimeter wavelengths, we may well have recovered the bulk of the

circumnuclear emission in  $^{12}\text{CO}$  (2–1). Our ATCA observation recovered  $76\% \pm 16\%$  of the HCN (1–0) emission detected by the SEST, suggesting once again that the bulk, if not all, of the circumnuclear emission in this line has been recovered. Surprisingly, our SMA observation recovered only  $26\% \pm 10\%$  of the  $^{13}\text{CO}$  (2–1) emission detected by the SEST, despite the fact that  $^{13}\text{CO}$  (2–1) traces higher column densities and therefore more compact regions than  $^{12}\text{CO}$  (2–1) as evidenced by the maps presented in Figure 8. The  $^{13}\text{CO}$  (2–1) line however was detected at a confidence level of only  $3\sigma$ – $4\sigma$  in the SEST observation, and we therefore do not place a high degree of confidence on this particular comparison.

##### 4.2. Physical Properties

We computed the physical properties (density and temperature) of the molecular gas associated with the circumnuclear starburst ring from the ratio in brightness temperatures of the HCN (1–0) to  $^{13}\text{CO}$  (2–1) lines using the large velocity gradient (LVG) approximation (e.g., Goldreich & Kwan 1974; Scoville & Solomon 1974). We used the  $^{13}\text{CO}$  (2–1) line because it is more likely than the  $^{12}\text{CO}$  (2–1) line to trace the same region as the HCN (1–0) line, as the maps we presented earlier indicate



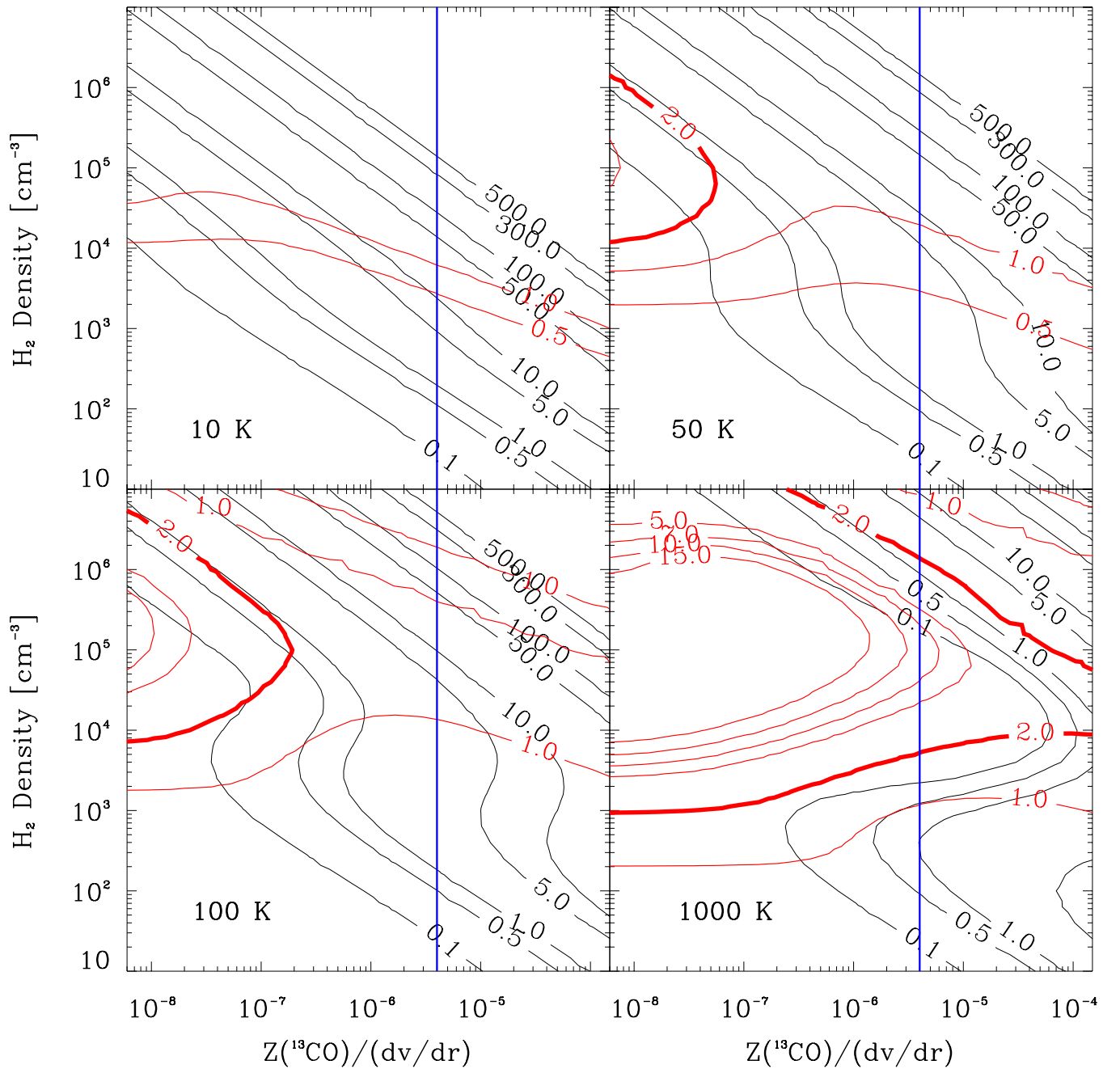
**Figure 11.** Upper left: integrated  $^{12}\text{CO}$  (2–1) intensity map in white/black contours plotted at 2, 5, 10, 15, 20, 25, 45, 65, and  $85 \times 4.5 \text{ Jy beam}^{-1} \text{ km s}^{-1}$  overlaid on the optical  $i'$ -band image in gray scale shown in Figure 1. Upper right: contours of the integrated HCN (1–0) intensity map (contour levels are the same as those shown in Figure 5) overlaid on the gray scale  $i'$ -band image. Lower left: integrated  $^{12}\text{CO}$  (2–1) (green contours) and HCN (1–0) (black contours) intensity maps both overlaid on the  $i'$ -band image (gray scale). Lower right: intensity-weighted mean  $^{12}\text{CO}$  (2–1) velocity in color scale, as well as white contours plotted in steps of  $20 \text{ km s}^{-1}$  starting from  $-60 \text{ km s}^{-1}$  and ending at  $60 \text{ km s}^{-1}$  with respect to the systemic velocity of  $1586 \text{ km s}^{-1}$ . In each panel, cyan contour is the outline of the  $5.8 \mu\text{m}$  emission at  $5.5 \text{ MJy sr}^{-1}$  from the *Spitzer* SINGS project published in Kennicutt et al. (2003); this emission traces the dust lane visible in silhouette in the optical image. A cross is plotted at the center of the galaxy, and the synthesized beam is shown at the lower left corner, of each panel.

(see Section 3.5 and Figure 8). HCN (1–0) is a tracer of high density molecular hydrogen gas; as the  $^{13}\text{CO}$  (2–1) emission is much more optically thin than the  $^{12}\text{CO}$  (2–1) emission, the former is more strongly weighted toward regions of high column density and therefore also density. Indeed, as we shall show, the  $^{13}\text{CO}$  (2–1) emission is optically thin and therefore a good tracer of column density.

To derive the ratio between the HCN (1–0) and  $^{13}\text{CO}$  (2–1) intensities, we first convolved our HCN (1–0) map to the same angular resolution as our  $^{13}\text{CO}$  (2–1) map (see Figure 8). The circumnuclear starburst ring is spatially unresolved at this angular resolution, and so the line ratio we derive for HCN (1–0) to  $^{13}\text{CO}$  (2–1) of  $2.0 \pm 0.2$  is an average over the entire ring. We assumed Galactic abundances for the  $^{13}\text{CO}$  and HCN molecules

of  $[\text{CO}]/[\text{H}_2] \approx 10^{-6}$  (Solomon et al. 1979) and  $[\text{HCN}]/[\text{H}_2] \approx 2 \times 10^{-8}$  (Irvine et al. 1987). Finally, we estimated a velocity gradient of  $0.24 \text{ km s}^{-1} \text{ pc}^{-1}$  from the P–V diagram along the kinematic major axis of the HCN (1–0) emission as shown earlier in Figure 9.

Figure 12 shows a plot of the density of molecular hydrogen gas,  $n(\text{H}_2)$ , as a function of the product  $Z(^{13}\text{CO})/(dv/dr)$ , where  $Z(^{13}\text{CO})$  is the  $^{13}\text{CO}$  abundance and  $dv/dr$  the velocity gradient along the radial direction. For the adopted or inferred values above,  $Z(^{13}\text{CO})/(dv/dr) \approx 4 \times 10^{-6}$ , which is displayed with the blue vertical line in each panel. The different panels in this figure correspond to different kinetic temperatures (from 10 K to 1000 K). The red contours in each panel indicate acceptable solutions in the  $n(\text{H}_2)$  and  $Z(^{13}\text{CO})/(dv/dr)$  phase

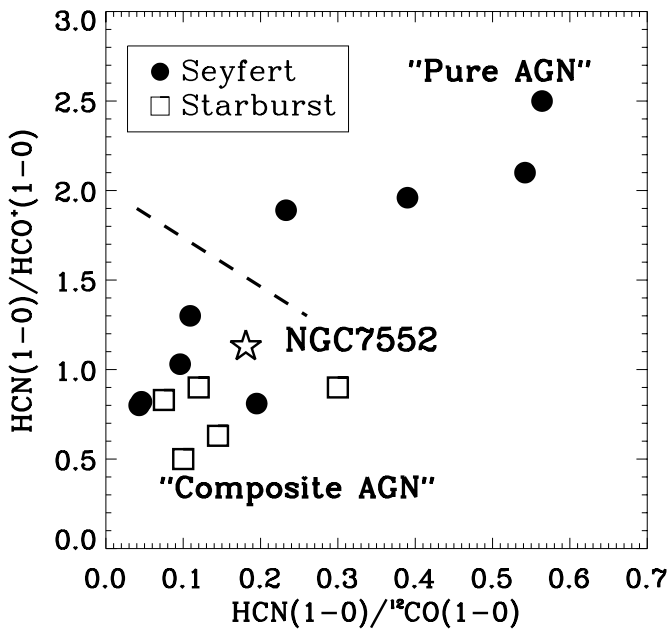


**Figure 12.** Derived physical properties of the molecular hydrogen gas in the circumnuclear starburst ring based on our LVG analysis. In all panels, the density of molecular hydrogen gas is plotted as a function of the product  $Z(^{13}\text{CO})/(dv/dr)$ , which for the present circumstances has an adopted value of  $\sim 4 \times 10^{-6}$  (vertical line in each panel; see text). Black contours in each panel indicate the  $^{13}\text{CO}$  (2-1) opacity. Red contours are the ratio in brightness temperature of HCN (1-0) to  $^{13}\text{CO}$  (2-1), which averaged over the ring has a value of  $2.0 \pm 0.2$  (thick red contour in each panel where visible). The upper left panel are solutions for a kinetic temperature of 10 K; upper right panel 50 K; lower left panel 100 K; and lower right panel 1000 K. As can be seen, for the measured line ratio and adopted product  $Z(^{13}\text{CO})/(dv/dr)$ , solutions can only be found for kinetic temperatures well above 100 K.

(A color version of this figure is available in the online journal.)

space for different HCN (1-0) to  $^{13}\text{CO}$  (2-1) line ratios based on our LVG calculations. The black lines indicate the  $^{13}\text{CO}$  (2-1) opacity. As can be seen, for a HCN (1-0) to  $^{13}\text{CO}$  (2-1) line ratio of  $2.0 \pm 0.2$ , and  $Z(^{13}\text{CO})/(dv/dr) \approx 4 \times 10^{-6}$ , solutions can only be found for kinetic temperatures well above 100 K, for which the  $^{13}\text{CO}$  (2-1) emission is optically thin. At such temperatures, the density of the molecular gas (averaged over the line-emitting regions) is around  $10^{5 \pm 1} \text{ cm}^{-3}$ . A similar result has been seen in the circumnuclear starburst ring of the barred spiral galaxy NGC 1097 (Hsieh et al. 2008).

On the other hand, the adopted Galactic abundance for  $^{13}\text{CO}$  of  $[^{13}\text{CO}]/[\text{H}_2] \approx 10^{-6}$  may be too high for starburst regions (Taniguchi & Ohya 1998). The  $^{13}\text{CO}$  abundance is controlled primarily by fractionation reactions and photodissociation. Fractionation reactions enhance the  $^{13}\text{CO}$  abundance, but such reactions are only effective in regions where the temperature is less than  $\sim 35$  K (Watson et al. 1976) and hence not exposed to vigorous star formation. Photodissociation, naturally, is even more effective in regions where the SFR is higher. If the combination of these factors reduces  $[^{13}\text{CO}]/[\text{H}_2]$  by, say, an



**Figure 13.** HCN (1–0)/<sup>12</sup>CO (1–0) vs. HCN (1–0)/HCO<sup>+</sup> (1–0) integrated intensity ratios of pure Seyfert and starburst galaxies. Solid circles indicate Seyfert galaxies, open squares starburst galaxies, and an open star the position of NGC 7552 (starburst) in this plot. Data other than for NGC 7552 are taken from Kohno (2005), Kohno et al. (2007, 2008), and Shibatsuka et al. (2003).

order of magnitude to  $\sim 10^{-7}$ , then acceptable solutions to the observed HCN (1–0) to <sup>13</sup>CO (2–1) line ratios can be found for temperatures  $\sim 50$ – $100$  K and densities  $\sim 10^5$  cm<sup>-3</sup>. Thus, unless the <sup>13</sup>CO abundance is even lower than the range of values considered above, the molecular gas that we detected in the starburst ring is relatively warm and dense, as would be expected for gas associated with regions that have recently undergone vigorous star formation (see also Matsushita et al. 1998, 2010).

In many galaxies exhibiting nuclear activity, it is often difficult to tell whether this activity is produced by active galactic nuclei (AGN) or starburst. Kohno et al. (2001) suggest that such galaxies can be distinguished by plotting the line ratio of HCN (1–0) to HCO<sup>+</sup> (1–0) against that of HCN (1–0) to <sup>12</sup>CO (1–0) (Figure 13). They show that galaxies dominated by or exhibiting only AGN activity lie in the upper right corner of such a diagram, having high values in both line ratios. By comparison, galaxies exhibiting only starburst activity or a mix of starburst and AGN activity lie in the lower left corner, having low values in both line ratios. In the case of NGC 7552, we can clearly see that the nuclear activity is dominated by a starburst (i.e., circumnuclear starburst ring), and our observations are able to confidently separate the molecular gas associated with the starburst from any associated with the nucleus in HCN (1–0) and HCO<sup>+</sup> (1–0), and also to some degree of confidence in <sup>12</sup>CO (2–1) and <sup>13</sup>CO (2–1) (based on the similarity of the gas kinematics in all four lines). Such is not the case for the galaxies used to make the abovementioned plot in Kohno et al. (2001). For NGC 7552, the HCN (1–0) to HCO<sup>+</sup> (1–0) line ratio, averaged over the circumnuclear starburst ring, is  $1.13 \pm 0.04$ . The HCN (1–0) to <sup>12</sup>CO (1–0) ratio, assuming that the <sup>12</sup>CO (1–0) is equal to the <sup>12</sup>CO (2–1) intensity for the reason mentioned earlier (see Section 4.1), is  $0.181 \pm 0.005$ . This places NGC 7552 firmly in the region occupied by galaxies with nuclear activity dominated by a starburst, enhancing the validity of the plot made by Kohno et al. (2001).

### 4.3. Dynamical Stability

We examine whether the molecular gas in the circumnuclear starburst ring is dynamically stable according to the Toomre criterion (Toomre 1964). For a disk to be stable against collapse due to its own gravity, the Toomre  $Q$  parameter, which is defined as the ratio of the critical surface gas density to the actual surface gas density,  $\Sigma_{\text{crit}}/\Sigma_{\text{gas}}$ , is required to have a value equal to or exceeding unity.

We compute the average surface gas density in a circular region of radius of  $10''$  ( $\sim 1$  kpc), just encompassing the circumnuclear starburst ring, using the molecular gas mass derived in Section 3.5 to find  $\Sigma_{\text{gas}} \approx 900 M_{\odot} \text{pc}^{-2}$ . Because the <sup>12</sup>CO (2–1) emission likely exhibits a central depression or hole within the circumnuclear starburst ring, the value we computed is likely to be a lower limit. We estimated the critical surface gas density using the relationship from Kohno et al. (1999)

$$\Sigma_{\text{crit}} = 630 \left( \frac{\sigma}{15 \text{ km s}^{-1}} \right) \left( \frac{\kappa}{0.57 \text{ km s}^{-1} \text{ pc}^{-1}} \right) M_{\odot} \text{pc}^{-2} \quad (1)$$

where  $\sigma$  is the one-dimensional velocity dispersion, for which we estimated a value of  $12 \text{ km s}^{-1}$  based on a second moment map of the <sup>12</sup>CO (2–1) emission (not shown), and  $\kappa$  the epicyclic frequency. As shown below in Section 5.2, we derived an epicyclic frequency of  $0.4 \text{ km s}^{-1} \text{ pc}^{-1}$  at the radius of the circumnuclear starburst, and therefore  $\Sigma_{\text{crit}} \approx 354 M_{\odot} \text{pc}^{-2}$ . The Toomre  $Q$  parameter is therefore no larger than 0.4 in the circumnuclear starburst ring, indicating that the molecular gas there is (highly) unstable to gravitational collapse.

## 5. GLOBAL KINEMATICS

In this section, we derive the global kinematics (rotation curve) of NGC 7552 and, assuming a pattern speed for the galactic bar, estimate the predicted locations of dynamical resonances. We then compare these predictions with the locations of prominent features in NGC 7552 that in other disk galaxies are commonly attributed (but seldom demonstrated explicitly to correspond) to dynamical resonances; in the case of NGC 7552, Forbes et al. (1994b) have highlighted features that possibly correspond to dynamical resonances. Our overarching goal is to determine whether we can understand the global kinematics of NGC 7552 sufficiently well so as to test ideas for the formation of dense molecular gas and stars in its circumnuclear starburst ring as discussed in the next section.

### 5.1. Rotation Curve

As pointed out in Section 3.2, the intensity-weighted mean H I velocity map (Figure 2) exhibits an S-shaped asymmetry characteristic of periodic orbits predicted theoretically for barred galaxies (e.g., Athanassoula 1992 and references therein). In general, such orbits are elliptical (and centered on the dynamical center of the galaxy), and have their major axes either parallel ( $x_1$  orbits) or orthogonal ( $x_2$  orbits) to the major axis of the bar, changing from one to the other between the locations of dynamical resonances. Specifically, beyond the ILR out to corotation (CR), orbits are  $x_1$ . Within the ILR, orbits are  $x_2$ . If there is both an outer ILR (oILR) as well as an inner ILR (iILR), orbits are  $x_1$  inside the iILR and  $x_2$  between the iILR and oILR. Between CR and the outer Lindblad resonance (OLR), orbits are  $x_1$ . Close to and beyond the major axis of the bar (predicted



theoretically to terminate at or before CR), both  $x_1$  and  $x_2$  orbits become progressively more (and are predicted to be usually close to) circular. Close to a bar potential,  $x_1$  is the predominant orbit, hence the major kinematic axis is closely aligned with the galactic bar. The change from  $x_1$  to  $x_2$  orbits and back between dynamical resonances causes the S-shaped twist in the major kinematic axis on scales up to that comparable with the length of the galactic bar.

To accurately derive the rotation curve of NGC 7552 from our measurements of the projected galaxy kinematics, we require a priori knowledge of the axial ratios (which change with galactic radius) of the  $x_1$  and  $x_2$  orbits, and the inclination of the galaxy (including any variation in inclination with galactic radius if the disk is warped). As these parameters (apart from perhaps inclination) are not in general known, we assumed a constant inclination of  $25^\circ$  for NGC 7552 based on the axial ratio of its stellar disk as described in Section 1. For the major kinematic axis of the disk, we used a P.A. of  $110^\circ$  as indicated by the first moment map in all four molecular lines as well as the inner region of the first moment map in the H I line.

The HCN (1–0) to HCO<sup>+</sup> (1–0) emissions originate primarily, if not entirely, from the circumnuclear starburst ring, and are therefore useful for inferring the rotational velocity of this ring. Sampling only a limited and furthermore spatially unresolved radial range, however, these lines do not provide a meaningful rotation curve at the center of the galaxy. The <sup>12</sup>CO (2–1) line samples a larger radial range from (presumably) the ring to the innermost regions of the dust lanes, and is useful for making a crude estimate of the rotation curve over this range. The H I line provides a relatively robust estimate of the rotation curve beyond the outer regions probed in <sup>12</sup>CO (2–1) to far beyond the stellar optical disk.

The rotation curve is directly derived from P–V diagrams of <sup>12</sup>CO (2–1) and H I from cuts along  $110^\circ$ , as shown earlier in Figure 9, with the envelope-tracing method (see details and more examples in Sofue 1996, 1997). The method traces terminal velocity (or tangent-points) on P–V diagram and corrects for the ISM velocity dispersion. The deprojected rotational velocity is derived from the terminal velocity  $V_t$ :

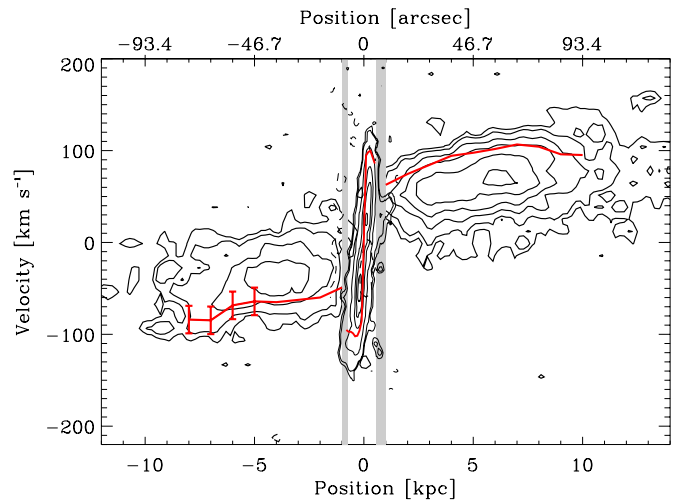
$$V_{\text{rot}} = V_t / \sin i - (\sigma_{\text{obs}}^2 + \sigma_{\text{ISM}}^2)^{1/2}, \quad (2)$$

where  $i$  is the inclination,  $\sigma_{\text{obs}}$  the velocity resolution of the observation, and  $\sigma_{\text{ISM}}$  the velocity dispersion of interstellar gas. We adopt an ISM velocity of  $7 \text{ km s}^{-1}$  (Stark & Brand 1989; Malhotra 1994). The terminal velocity is the velocity at which the intensity ( $I_t$ ) is equal to

$$I_t = [(\eta I_{\text{max}})^2 + I_{\text{lc}}^2]^{1/2} \quad (3)$$

on the P–V diagram, where  $I_{\text{max}}$  and  $I_{\text{lc}}$  are, respectively, the maximum intensity and intensity of the lowest (typically the  $3\sigma$ ) contour level. If the intensity is high enough,  $I_t \sim I_{\text{lc}}$  and is therefore approximately defined by the locus of the rotation curve.  $\eta$  is taken as 0.2 as suggested in Sofue (1996), corresponding to the 20% level of the intensity profile at a fixed position, i.e.,  $I_t \simeq 0.2 \times I_{\text{max}}$ . The accuracy of the obtained  $V_t$  along with the  $V_{\text{rot}}$  by the envelope-tracing method is  $\sim \pm 15 \text{ km s}^{-1}$ , corresponding to  $\sim \pm 15 \text{ km s}^{-1} / \sin i$  in deprojected velocity (Sofue 1996, 1997).

In Figure 14, we plot the projected rotation curve derived from the <sup>12</sup>CO (2–1) (from 0 to  $\sim 0.8 \text{ kpc}$ ) and H I (beyond  $1 \text{ kpc}$ ) lines onto the P–V diagram of both these lines cut along a P.A. of  $110^\circ$ . The projected uncertainties of  $\pm 15 \text{ km s}^{-1}$  are plotted



**Figure 14.** Rotation curve, projected for an inclination of  $25^\circ$ , shown in red curves overlaid on a P–V diagram derived from our <sup>12</sup>CO (2–1) ( $-1$  to  $+1 \text{ kpc}$ ) and H I maps ( $\leq -1$  and  $\geq 1 \text{ kpc}$ ) at a P.A. of  $110^\circ$  shown in contours. Shaded regions indicate locations where the two tracers connect, and at which the rotation curve fail to track (see text). The typical projected uncertainty of the rotation curve obtained by the envelope tracing method is  $\pm 15 \text{ km s}^{-1}$ , and illustrated at the end of the approaching side. Contours in <sup>12</sup>CO (2–1) are plotted in steps of  $-2\sigma$ ,  $2\sigma$ ,  $3\sigma$ ,  $5\sigma$ ,  $10\sigma$ ,  $20\sigma$ ,  $30\sigma$ ,  $40\sigma$ , and  $45\sigma$ , where  $1\sigma = 78 \text{ mJy beam}^{-1}$ , and contours in H I diagram in steps of  $-2\sigma$ ,  $2\sigma$ ,  $3\sigma$ ,  $5\sigma$ ,  $7\sigma$ ,  $10\sigma$ ,  $15\sigma$ , and  $20\sigma$ , where  $1\sigma = 1.1 \text{ mJy beam}^{-1}$ .

(A color version of this figure is available in the online journal.)

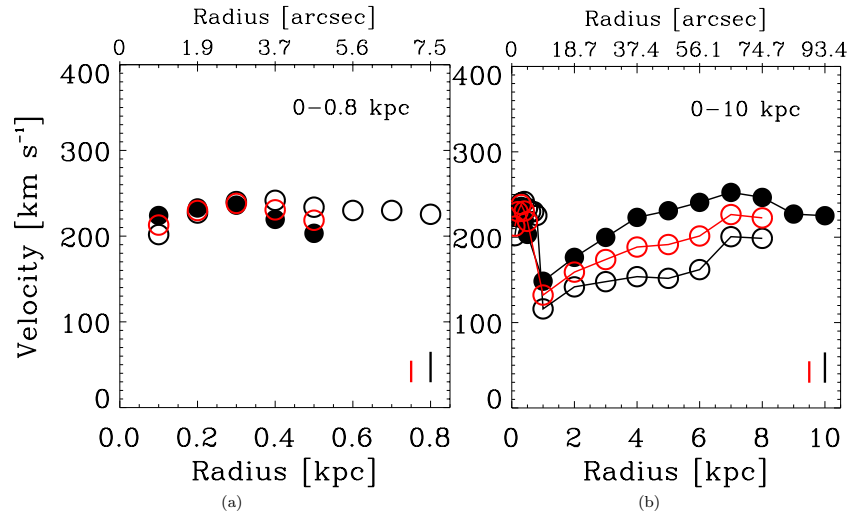
on the approaching side of the rotation curve to indicate the size of the error bars. Due to the imperfect physical connection between the two tracers and asymmetry of P–V diagrams, the rotation curve between  $0.5$  to  $1 \text{ kpc}$  is poorly connected; we indicate these regions with gray shading in Figure 14.

The deprojected inner and the entire derived rotation curves are shown respectively in Figures 15(a) and (b). The open and solid black circles denote the receding and approaching sides of the rotation curve respectively, and the red open circles indicate the average of the two sides at the radii where both receding and approaching velocities are seen. The deprojected uncertainties are plotted in the lower right of each panel.

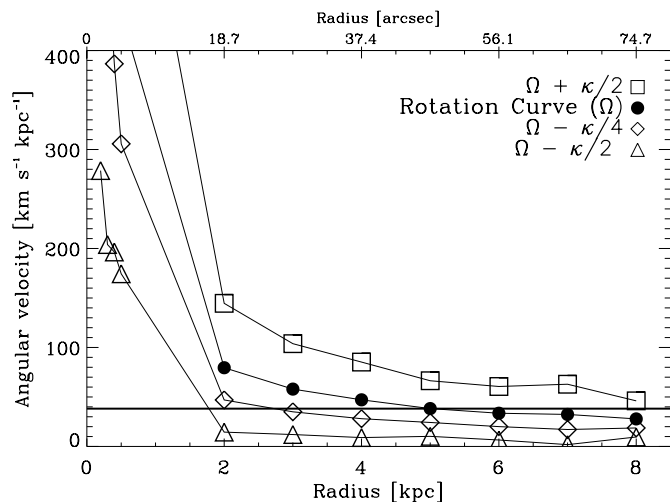
## 5.2. Dynamical Resonances

Because the intensity of the approaching side in the H I map is significantly weaker than that of the receding side, together with the asymmetry of the derived rotation curve on each side, we determine the dynamical resonances for each side of the galaxy separately. We compute where dynamical resonances are theoretically predicted to occur using the usual assumption that the angular velocity of the bar (its pattern speed,  $\Omega_p$ ) is equal to the angular velocity of the galaxy at the outermost ends of the bar (i.e., the CR). We determined the length of the bar in NGC 7552 from the Two Micron All Sky Survey (2MASS) image at  $2.2 \mu\text{m}$ , assuming that the bar ends where the two prominent spiral arms first appear to emerge at a radius of about  $5 \text{ kpc}$ .

The average rotation curve shows that the angular velocity of the galaxy at the bar’s end ( $5 \text{ kpc}$ ) is about  $38 \text{ km s}^{-1} \text{ kpc}^{-1}$ , which we therefore take to be the pattern speed of the bar. In Figure 16, we plot the angular velocity (velocity divided by radius,  $V/R$ ) of the galaxy,  $\Omega$  (filled circles), as derived from the rotation curve. The solid line is an interpolation between adjacent angular velocity data points, from which we derived the epicyclic frequency,  $\kappa$ , of the orbits, where



**Figure 15.** Deprojected rotation curves derived for NGC 7552 from our  $^{12}\text{CO}$  (2–1) maps ( $<1$  kpc) and H I maps ( $\geq 1$  kpc). Left panel shows the rotation curve within 1 kpc, and right panel the entire rotation curve up to 10 kpc. Solid and open black circles denote the rotation curve of the receding and approaching sides, respectively. We assumed a constant inclination of  $25^\circ$  and P.A. of  $110^\circ$  for the kinematic major axis. Red circles show the average rotation curve at the radii where both the receding and approaching sides can be measured. The uncertainties of the individual and average rotation curves are indicated in the lower right corner of each panel. (A color version of this figure is available in the online journal.)



**Figure 16.** Diagram used to infer the locations of dynamical resonances. Black solid circles indicate the angular velocity of the galaxy,  $\Omega$ , based on an interpolation of the rotation curve shown in Figure 15. Individual measurements that define the  $\Omega - \kappa/2$  (ILR) curve are plotted as triangles, the  $\Omega - \kappa/4$  (UHR) curve as diamonds, and the  $\Omega + \kappa/2$  (OLR) curve as squares, where  $\kappa$  is the epicyclic frequency of the orbit. Resonances occur where the values of these curves are equal to the pattern speed of the bar, indicated by the thick horizontal line.

$\kappa^2 = 2((v^2/R^2) + (v/R)(dv/dR))$ . We then computed the  $\Omega - \kappa/2$  curve (triangles) to infer the location(s) of the ILRs, the  $\Omega - \kappa/4$  curve (diamonds) to infer the location of the Ultra Harmonic Resonance (UHR), and the  $\Omega + \kappa/2$  curve (squares) to infer the location of the OLR. The uncertainties in these curves are rather large so we do not show the error bars in Figure 16 for clarity, but show the possible radial ranges of the resonances taking these uncertainties into account.

We therefore constrain the following possible locations of dynamical resonances in NGC 7552. The  $\Omega + \kappa/2$  curve (open squares) cross the value of  $\Omega_p$  at  $>8$  kpc, the predicted location of the OLR. Theory predicts that the OLR in spiral galaxies should occur just beyond the outermost reaches of their spiral arms, which in the case of NGC 7552 can be traced out to a

(deprojected) radius of about 10 kpc. The computed  $\Omega - \kappa/4$  curve implies that the UHR occurs at roughly 2–4 kpc. Forbes et al. (1994b) suggest that an isolated bright H II region located at a (deprojected) radius of 3.5 kpc coincides with the UHR. The  $\Omega - \kappa/2$  curve implies two ILRs. The oILR is located at a radius of about 1.7 kpc, which is beyond the circumnuclear starburst ring that has a radius of about 0.5 kpc. Further inward the angular velocity might eventually drop to a value of  $0 \text{ km s}^{-1}$  (in principle) at the center of the galaxy, and so an iILR could exist somewhere within the circumnuclear starburst ring. On the other hand, we do not exclude the possibility that the rotation curve keeps rising due to a black hole in the galactic center, resulting in the absence of an iILR. Our finding that the circumnuclear starburst ring is located within the oILR and beyond, possibly, the iILR, is consistent with theoretical predictions (Buta & Combes 1996). The reasonable agreement between the computed and anticipated locations of dynamical resonances in NGC 7552 provides a measure of confidence in our understanding of the global kinematics of this galaxy.

## 6. STAR FORMATION IN THE CIRCUMNUCLEAR RING

In this section, we investigate the star formation properties of the circumnuclear starburst ring, determine whether the observed properties of the ring are consistent with the model proposed by Kohno et al. (1999), and if not what modifications are required.

### 6.1. Star Formation Rate

The radio continuum observations of Forbes et al. (1994a), the data from which we reprocessed here, provides the cleanest (i.e., not subject to extinction, emission detected at high S/N) measure of the SFR in the circumnuclear ring. Like Forbes et al. (1994a), we used the standard prescription suggested by Condon (1992) to derive the supernova rate as  $\sim 0.1 \text{ yr}^{-1}$  from the measured radio luminosity for the ring. We then estimated the SFR from the supernova rate in the manner also suggested by Condon (1992), but corrected for a Salpeter initial mass function (IMF) as suggested by Condon et al.

(2002), as about  $15 M_{\odot} \text{ yr}^{-1}$ . An independent estimate of the SFR in the ring can be made from hydrogen recombination lines in the optical and near-infrared, although such estimates are affected by extinction. The luminosity of the  $H\alpha$  line of  $1.5 \times 10^{42} \text{ erg s}^{-1}$  derived from the Gemini image, using the relationship  $L_{H\alpha}$ –SFR suggested by Kennicutt (1998), implies an SFR of about  $12 M_{\odot} \text{ yr}^{-1}$ . Moorwood & Oliva (1988) find a luminosity in the  $\text{Br}\gamma$  line of  $9.8 \times 10^{39} \text{ erg s}^{-1}$  for the circumnuclear ring. Using the relationship between the  $\text{Br}\gamma$  line luminosity and SFR suggested by Kennicutt (1998), the SFR is about  $8 M_{\odot} \text{ yr}^{-1}$ . The correlation between the global HCN (1–0) and far-infrared luminosities found by Gao & Solomon (2004) for normal to extreme star-forming galaxies can be used to provide an estimate of the SFR using the measured HCN (1–0) luminosity ( $L'_{\text{HCN}}$ ) as a proxy for the infrared luminosity ( $L_{\text{IR}}$ ). With an  $L'_{\text{HCN}} \approx 7.5 \times 10^7 \text{ K km s}^{-1} \text{ pc}^2$  and hence  $L_{\text{IR}}$  of  $(5.5 \pm 0.1) \times 10^{10} L_{\odot}$  in the circumnuclear starburst ring, the corresponding SFR (for a Salpeter IMF) is  $11.0 \pm 0.2 M_{\odot} \text{ yr}^{-1}$ . Therefore, in summary, the SFR in the circumnuclear starburst ring of NGC 7552 is about  $10$ – $15 M_{\odot} \text{ yr}^{-1}$ .

For comparison, the global SFR derived in the same manner from observations of the entire galaxy in the radio continuum at 20 cm by Schmitt et al. (2006) is about  $18 M_{\odot} \text{ yr}^{-1}$ . The global SFR derived from the  $H\alpha$  line, based on our Gemini observations, is about  $47 M_{\odot} \text{ yr}^{-1}$ . The global SFR can also be estimated from the far-IR measurements of NGC 7552 with *IRAS* by Sanders et al. (2003). They find an infrared luminosity for this galaxy of  $1.1 \times 10^{11} L_{\odot}$ , thus placing NGC 7552 right at the threshold for a luminous infrared galaxy. Using as above the conversion between infrared luminosity and SFR suggested by Gao & Solomon (2004) for a Salpeter IMF, the SFR is about  $22 M_{\odot} \text{ yr}^{-1}$ . Because the SFR estimated from  $H\alpha$  can be highly uncertain (e.g., up to 99% of UV photons can leak out of H II regions; Kennicutt 1998), we give more weight to the other methods of estimating SFR as described above. Thus, the circumnuclear region of NGC 7552 contributes  $\sim 50\%$ – $80\%$  of the galaxy’s entire star formation, and being strongly concentrated in the circumnuclear ring, results in an extremely high SFR surface density of  $\gtrsim 4 \times 10^{-6} M_{\odot} \text{ yr}^{-1} \text{ pc}^{-1}$ .

The SFR in the circumnuclear ring, divided by the mass of molecular gas present, is about  $(3$ – $5) \times 10^{-9} \text{ yr}^{-1}$ . The inverse of this ratio represents the timescale for gas exhaustion and is  $\sim 3 \times 10^8 \text{ yr}$ . Such a short gas exhaustion timescale is comparable to that of interacting objects and ULIRGs (few  $\times 10^8 \text{ yr}$ ) and about 10 times shorter than that in normal galaxies (Bigiel et al. 2008). Note that even though the gas exhaustion timescale of a circumnuclear starburst ring is short, as the galaxies with circumnuclear rings are generally harboring stronger bars than the galaxies lacking rings, the gas is likely to be nearly continuously replenished at the circumnuclear region through, presumably, the bar-driven gas inflow.

### 6.2. Formation of Dense Molecular Gas and Stars at the Circumnuclear Ring

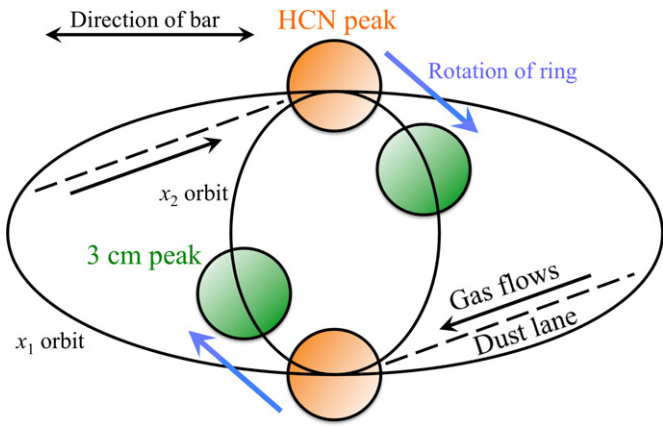
The most prominent optical features—circumnuclear starburst ring, large-scale bar, and spiral arms—in NGC 7552 are remarkably similar to those of NGC 6951, including even the orientation of these features on the sky, thus providing a simple and direct comparison. Like NGC 7552, the bar in NGC 6951 is oriented east–west, with spiral arms emerging from the bar and curling north on the western side and south on the eastern side. Dust lanes that run along the bar intersect the circumnuclear

starburst ring at its northern and southern sides, just like those seen in NGC 7552.

Kohno et al. (1999) reported observations of the circumnuclear starburst ring of NGC 6951 in both  $^{12}\text{CO}$  (1–0) and HCN (1–0) with resolution of  $\sim 500 \text{ pc}$  in both lines. They found twin CO knots at or close to the location where the dust lanes meet the ring, and twin HCN knots displaced downstream (assuming that the spiral arms are trailing) from the CO knots. On this basis, Kohno et al. (1999) suggested that the molecular gas at the locations where the dust lanes meet the ring—the twin CO knots, comprising gas channeled inward by the bar—is too turbulent to form dense molecular clouds. Instead, it is only when this gas moves from  $x_1$  orbits (closely traced by the dust lanes) beyond the oILR to  $x_2$  orbits within the oILR (the circumnuclear starburst ring) that dense molecular gas—seen as the twin HCN knots—are able to form; the velocity dispersion along  $x_2$  orbits is predicted to be much smaller than that along  $x_1$  orbits, and hence the molecular gas along  $x_2$  orbits is much more susceptible to gravitational instabilities (see their Figure 17). This scenario is not supported in our observation at least in a global sense as we find that the molecular gas in the circumnuclear starburst ring is very unstable to collapse.

Krips et al. (2007) have observed the circumnuclear starburst ring of NGC 6951 at significantly higher angular resolution ( $\sim 240 \text{ pc}$ ) and sensitivity also in  $^{12}\text{CO}$  (2–1) and HCN (1–0). An examination of their results (see their Figure 1) reveals no discernible displacement between the centroids of the CO knot and dominant HCN knot in the northern part of the ring, and at best a small displacement between the centroids of the CO knot and dominant HCN knot in the southern part of the ring. (There is a weaker HCN knot located just downstream of each dominant HCN knot: together, these knots, when convolved to the lower angular resolution attained in the observation by Kohno et al. 1999, produce an apparent angular displacement between the HCN and CO knots.) The centroids of all these knots are located at or close to where the dust lanes meet the ring, just as we see for the two HCN knots in the circumnuclear starburst ring of NGC 7552. It is possible that collisions between gas clouds at the intersection of  $x_1$  orbits (closely traced by the dust lanes) and the  $x_2$  orbits (the circumnuclear starburst ring) give rise to relatively dense molecular gas (e.g., Combes & Gerin 1985), although the relative collision velocity must not be so high ( $\lesssim 50 \text{ km s}^{-1}$ ) as to dissociate or ionize the molecular gas. Alternatively, or in addition, turbulence generated by shocks help compress the molecular gas (Elmegreen 1994); turbulence can help support GMCs against collapse on large scales, and promote the formation of dense condensations—from which stars form—on small scales.

As sketched in Figure 17, both our observations of NGC 7552 and that by Krips et al. (2007) of NGC 6951 therefore suggest an important revision to the picture proposed by Kohno et al. (1999). In both these galaxies, we suggest that collisions between diffuse (CO) molecular gas being channeled inward along the dust lanes, turbulence generated by shocks, along with gravitational instabilities perhaps help promote the formation of dense molecular gas at or close to the intersection of the  $x_1$  and  $x_2$  orbits where the dust lanes meet the circumnuclear starburst ring. Indeed, for NGC 7552, the ratio in brightness temperature of the HCN (1–0) to  $^{12}\text{CO}$  (2–1) emission averaged over the ring is  $\sim 0.18$ , whereas that at the stronger  $^{12}\text{CO}$  (2–1) western arm has a  $3\sigma$  upper limit of  $\sim 0.12$ . This indicates that the formation of dense molecular gas is promoted where the dust lanes meet the ring, relative to regions immediately



**Figure 17.** A sketch showing our zeroth-order picture of the formation of dense gas (HCN peaks) and synchrotron emission (3 cm peaks) resulting from supernova explosions of stars that formed from the dense gas in the circumnuclear starburst ring of NGC 7552. This sketch is based on a revision of the picture proposed by Kohno et al. (1999) for NGC 6951.

(A color version of this figure is available in the online journal.)

beyond, along the same (at the very least for the western) dust lanes.

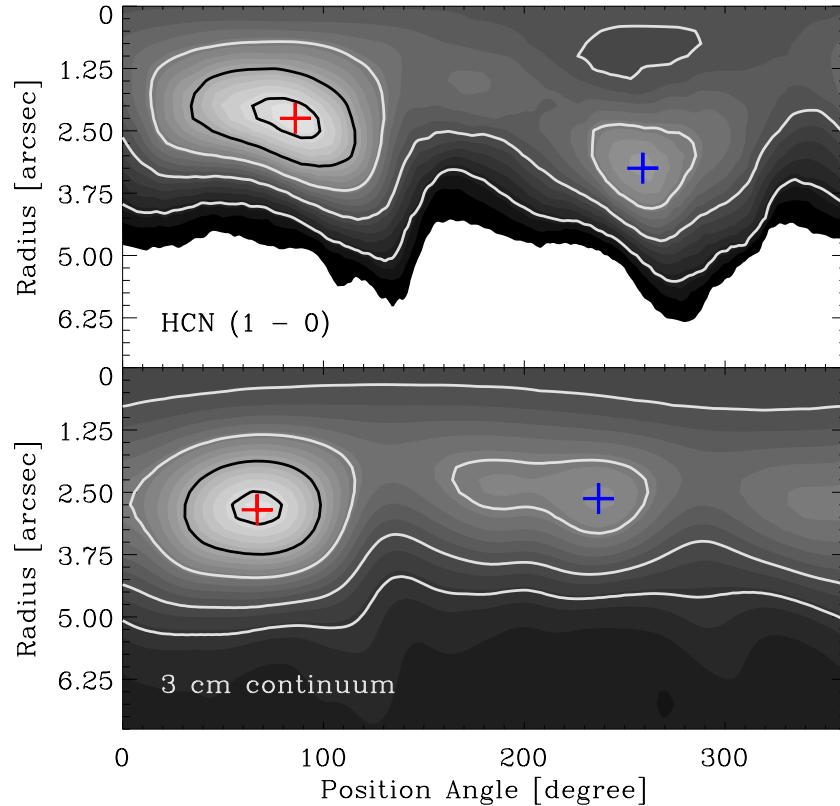
Our revised picture also helps explain another observed property of the circumnuclear starburst ring in both NGC 7552 and NGC 6951, and that is the displacement between the centroids of their HCN and radio continuum knots. Such a displacement was pointed out above for NGC 7552 (see Section 3.4 and Figure 5), and is obvious even through a cursory visual comparison of the molecular line maps made by Krips et al. (2007) with the radio continuum maps made by Saikia et al. (2002) in NGC 6951. If we assume that the spiral arms in NGC 7552 are trailing as assumed by Kohno et al. (1999) for NGC 6951, then the radio continuum knots are located downstream of the HCN knots in both NGC 7552 and NGC 6951. The most straightforward explanation for this displacement is that it reflects the timescale between the formation of massive stars from dense molecular gas at or close to the intersection of the  $x_1$  and  $x_2$  orbits, and the eventual demise of these stars in supernova explosions. This argument can be checked by calculating the time it would take for the rotation of the circumnuclear starburst ring to have carried the radio continuum knots downstream away from the present location where the dust lanes meet the ring (specifically, the present location of the HCN knots), keeping in mind that during this time the pattern speed of the bar also carries the dust lane downstream with respect to the ring, albeit by a smaller angular distance.

Figure 18 shows the 3 cm radio continuum map of Figure 3 and HCN integrated intensity map of Figure 5 plotted in polar coordinates. The northern radio and HCN knots have a deprojected angular offset  $0''.47 \pm 0''.05$  ( $\sim 50$  pc), and the southern radio and HCN knots a deprojected angular offset of  $0''.67 \pm 0''.08$  ( $\sim 70$  pc). As shown in Section 5.1, the average rotational velocity of both sides of the circumnuclear starburst ring is about  $218 \text{ km s}^{-1}$ , and the pattern speed of the bar about  $38 \text{ km s}^{-1} \text{ kpc}^{-1}$ , corresponding to  $19 \text{ km s}^{-1}$  at the ring (0.5 kpc); their net velocity is therefore about  $199 \text{ km s}^{-1}$ . The propagation time between the northern radio and HCN knots is therefore  $(2.4 \pm 0.3) \times 10^5$  year, and that between the southern radio and HCN knots  $(3.4 \pm 0.4) \times 10^5$  year. These timescales are about an order of magnitude shorter than the 5–10 Myr timescale for OB stars to explode as supernovae. It is possible

that the true pattern speed is significantly higher and/or the rotational velocity of the ring significantly slower than what we inferred, although it seems unlikely that both these speeds have been underestimated by up to an order of magnitude.

In the case of NGC 6951, the centroids of the twin radio continuum knots imaged by Saikia et al. (2002) coincide closely with the centroids of the twin HCN knots as imaged by Kohno et al. (1999), who estimated a propagation time between their observed twin CO and HCN knots of a few million years. This is in better agreement with, although as in NGC 7552 on the shorter side of, the timescale for the demise of massive stars with a Salpeter-like IMF. We point out that the model proposed above for the formation of dense gas and stars in the circumnuclear starburst ring can be more cleanly applied to NGC 6951 than NGC 7552. The latter exhibits not just two but five radio continuum knots in its circumnuclear starburst ring (see Figure 3, right panel), indicating that the actual pattern for star formation—and by extension formation of dense molecular gas—is more complicated than can be fully captured by the simple model proposed in Figure 17. In reality, dense molecular gas may form in the dust lanes, as they have been found in some bar galaxies with HCN (1–0) observations (e.g., Leon et al. 2008), before being channeled into the circumnuclear starburst ring (albeit, as shown in the case of NGC 7552, not as efficiently as at the intersection of these dust lanes with the ring), and continue to form in the circumnuclear starburst ring downstream from where the dust lanes meet the ring (e.g., through gravitational instabilities).

Even in the simpler case of NGC 6951, let alone NGC 7552, there are further complications that need to be taken into account. The formation of dense molecular gas at any given location may not proceed at a constant rate but is modulated in time. Some star formation may therefore occur even before gas is channeled into the ring and continue further downstream away from where the dust lanes meet the ring. In addition, we are not able to separate thermal and non-thermal contributions to the radio continuum at this stage, and thus cannot attribute all the emission to a post-supernova stage. Condon (1992) suggests a simple way to separate the thermal emission from the total observed radio flux at a given frequency based on the spectral index. For a spectral index of  $-0.89$  derived in Section 3.3, the fraction of thermal emission is about 40% at 3 cm ( $\sim 10$  GHz). If thermal (i.e., free-free emission from H II regions) is enhanced at or just downstream of the HCN knots due to star formation, then the convolution (due to the limited angular resolution of our observation) of this emission with non-thermal emission produced further downstream by supernovae would tend to shift the centroid of the combined emission upstream toward the HCN knots. The result is a smaller apparent propagation time between the HCN and radio continuum knots than is actually the case. On the other hand, the bright millimeter lines seen at the contact points may reflect high velocity dispersions (Downes et al. 1993) and therefore large line widths, and the concentrations of dense molecular gas actually peak somewhat downstream of the HCN knots. In this case, the result is a larger apparent propagation time between the HCN and radio continuum knots than is actually the case. Finally, the IMF for star formation in the ring may be top-heavy, resulting in a shorter than usual timescale between the formation and death of the majority of massive stars (specifically, the time at which the radio intensity of all their supernovae combined peaks following the formation of these stars at the same epoch) compared with a Salpeter-like IMF. Papadopoulos (2010) and Papadopoulos et al. (2011) have



**Figure 18.** Integrated HCN (1–0) intensity of Figure 5 (upper panel) and 3 cm continuum intensity of Figure 3 (lower panel) plotted in polar coordinates. The abscissa is polar angle measured from east, and the ordinate radius from center. Crosses in both panels indicate the centroids of prominent pairs of knots in the emission.

(A color version of this figure is available in the online journal.)

pointed out that, in regions of high star formation densities, the resulting high cosmic-ray densities can heat molecular gas to much higher temperatures than is the case in Galactic GMCs, in agreement with the relatively high temperatures than we infer for the molecular gas in the starburst ring of NGC 7552. The higher temperature of the molecular gas requires a correspondingly higher characteristic Jeans mass for star formation in this gas, resulting in (1) stars with larger masses than are found in Galactic GMCs and (2) a top-heavy IMF as low-mass star formation is suppressed (Papadopoulos et al. 2011).

Our revised model of the sequence of star formation in this circumnuclear ring is consistent with the *pearls on a string* scenario of Böker et al. (2008, see their Figure 7), which is based on their near infrared (*H*- and *K*-band) observations along with the age identification using the intensity ratios of emission lines. The *pearls on a string* model proposes that the overdense region is close to where the gas enters the ring, and therefore where the starburst is triggered. The star clusters formed subsequently will continue their orbit along the circumnuclear ring. In this scenario, a bipolar age gradient is expected, as has been seen in some galaxies, based on infrared and optical studies, such as NGC 1068 (Davies et al. 1998), NGC 7771 (Smith et al. 1999), M100 (Ryder et al. 2001; Allard et al. 2005), IC 4933 (Ryder et al. 2010), and M83 (Knäpen et al. 2010).

The infrared age dating method has been applied to NGC 7552 by Brandl et al. (2012), who identify nine unresolved mid-infrared peaks (star clusters) in the starburst ring. There are tentative signs of an age gradient, in that the relatively young clusters are located near the north contact point, while the older clusters are found more than a quarter of the way

around the ring away from the contact points. This result is qualitatively consistent with the age gradient expected by the *pearls on a string* scenario and our revised model, although it must be noted that the age difference among the clusters is rather small at between only 5.5 and 6.3 Myr (see Figure 10 in Brandl et al. 2012).

Even though the aforementioned bipolar age gradients are found in the circumnuclear ring surveys of Böker et al. (2008) and Mazzuca et al. (2008), the detection rate of such age gradient patterns is only 60% and 43%, respectively, in these two surveys. We cannot therefore rule out the possibility of other scenarios for star formation, in addition to that described above, in the circumnuclear starburst ring of NGC 7552. For example, Böker et al. (2008) suggest a second *Popcorn* model in which the gas clouds become gravitationally unstable due to turbulence (Elmegreen 1994), and collapse at more random times and locations. Accordingly, there is no clear age gradient expected in the *Popcorn* model, as found by Hsieh et al. (2011) for the circumnuclear starburst ring of NGC 1097.

In NGC 7552, we detect five radio continuum knots, not just two. In NGC 6591, apart from the two dominant HCN(1–0) peaks, there are two other HCN(1–0) knots in the circumnuclear starburst ring. Thus, even for NGC 7552 and NGC 6591, the overall picture for the formation of dense molecular gas and stars in circumnuclear starburst rings cannot be as simple as that sketched in Figure 17. In reality, the formation of dense molecular gas may be promoted at the contact points, but also elsewhere in the circumnuclear starburst ring and star formation also occur there. Both the *pearls on a string* and *Popcorn* models may operate simultaneously in any given circumnuclear starburst ring, complicating any single interpretation. What we

present here is a zeroth-order model; the true picture requires observations at higher angular resolutions to properly study all locations where dense molecular gas is forming.

## 7. SUMMARY AND CONCLUSIONS

NGC 7552 is a barred galaxy that, like a significant fraction ( $\sim 20\%$ – $50\%$ ; Knapen 2005) of such galaxies, exhibits a circumnuclear starburst ring. A pair of dust lanes can be seen running along the bar on opposite sides of the center of the galaxy, and at their inner ends connect with the northern and southern parts of the circumnuclear starburst ring. The circumnuclear ring currently contributes 50%–80% of the galaxy’s entire star formation, and has a gas consumption timescale intermediate between normal and extreme star-forming (ultraluminous-infrared) galaxies.

To better understand the sequence and circumstances under which dense molecular gas and stars form in this ring, we have observed the center of NGC 7552 in molecular lines tracing relatively diffuse molecular gas ( $^{12}\text{CO}$  and  $^{13}\text{CO } J = 2-1$ ), which is the characteristic of Galactic GMCs; and relatively dense molecular gas ( $\text{HCN}$  and  $\text{HCO}^+ J = 1-0$ ), which is the characteristic of dense condensations in GMCs where stars form. We find that emission from (1) the dense molecular gas tracers coincide with the circumnuclear starburst ring, and exhibit two peaks located at or close to where the dust lanes meet this ring, and (2) the diffuse molecular gas tracers also coincide with the circumnuclear starburst ring, and  $^{12}\text{CO } (2-1)$  also with the innermost ends of the two dust lanes.

Using the LVG approximation, we demonstrated the following.

1. The molecular gas as observed averaged over the circumnuclear starburst ring is both relatively dense ( $10^5 \pm 1 \text{ cm}^{-3}$ ) and warm (well over 100 K), suggesting that massive stars have indeed formed from this gas.
2. The diffuse molecular gas in the circumnuclear starburst ring is unstable to gravitational collapse, providing at least one pathway for the formation of dense molecular gas.

We recreated maps of the H I emission from NGC 7552 using the data taken by Dahlem (2005), but at a higher angular resolution than those previously published. We used this map, along with our maps in molecular lines, to derive an approximate rotation curve for NGC 7552, and computed the predicted locations of dynamical resonances under the usual assumption that the outer ends of the bar terminate at the CR. We found that, within the statistical uncertainties, our results (1) are consistent with the idea that the OLR coincides closely with the outermost reaches of the twin spiral arms (which emerge from the ends of the bar) in NGC 7552, as theory predicts, and (2) show that the circumnuclear starburst ring likely lies between the iILR and oILR, as theory also predicts.

We then examined what our results imply for the formation of dense molecular gas and stars in the circumnuclear starburst ring of NGC 7552, specifically in the context of the picture proposed by Kohno et al. (1999) based on their observations of NGC 6951. The results and conclusions are as follows.

1. In both NGC 7552 and NGC 6951, emission from the dense molecular-gas tracers peaks at or close to the location where the pair of dust lanes that run along the bar meet the circumnuclear starburst ring (in the case of NGC 6951, as later improved observations by Krips et al. 2007 also show). This result suggests that the formation of dense molecular

gas is promoted where the  $x_1$  orbits (as closely traced by the dust lanes) and  $x_2$  orbits (circumnuclear starburst ring) intersect and where collisions between gas clouds is therefore more likely, rather than farther downstream along the ring as proposed by Kohno et al. (1999) where the velocity dispersion of  $x_2$  orbits is predicted to be much smaller than that of  $x_1$  orbits and hence the gas more susceptible to gravitational instabilities. Indeed, we observe that the formation of dense molecular gas is promoted where the dust lanes meet the ring compared with regions immediately beyond along the same dust lanes.

2. Our revised model as sketched in Figure 17 also naturally explains the observed displacement between the centroids of the dense molecular gas and radio continuum knots in both NGC 7552 and NGC 6951. This displacement ought to reflect the timescale between the formation of massive stars in the dense gas located where the dust lanes meet the circumnuclear starburst ring, and the demise of these stars as supernovae. In the case of NGC 6951, the estimated timescale between the centroids of these knots is a few million years, at the shorter end of the 5–10 Myr timescale for the majority of massive stars formed with a Salpeter IMF to become supernovae, and in the case of NGC 7552 is an order of magnitude shorter still. This may reflect uncertainties in the inferred angular velocities of the ring and bar, together with additional complexities mentioned below or a top-heavy IMF.

Finally, we noted that the observed properties of the circumnuclear starburst ring in NGC 7552 are more complicated than can be fully explained by the simple model proposed earlier and sketched in Figure 17. This model can be more cleanly applied to NGC 6951, whereas the circumnuclear starburst ring of NGC 7552 exhibits not just two, but five, knots in radio continuum. Our observations do not have sufficiently high angular resolution to spatially separate these knots in the dense (or diffuse) molecular gas tracers, so that the picture that emerges from our analyses represents a zeroth-order model for the formation of dense gas and stars in the circumnuclear starburst ring of NGC 7552. Observations at higher angular resolution and sensitivity are necessary to tease out details not incorporated into our model, which itself is based on an important revision of the model originally proposed by Kohno et al. (1999) for NGC 6951.

S.M. acknowledges grants from the National Science Council of Taiwan, NSC 100-2112-M-001-006-MY3, NSC 97-2112-M-001-021-MY3 to support Hsi-An Pan as a Masters student, and for support to conduct this work. T.W. acknowledges support from an ARC-CSIRO Linkage Fellowship in the early stages of this project and from the University of Illinois. The winning entry in the 2011 Australian Gemini School Astronomy Contest was submitted by Benjamin Reynolds, a student at Sutherland Shire Christian School. We are grateful to Benjamin, Chris Onken from the Australian Gemini Office, and Dr. Travis Rector from the University of Alaska, Anchorage, for making these data available to us. We thank Michael Dahlem for furnishing the reprocessed ATCA 3 cm, 6 cm, and 21 cm images. We thank Dr. Kotaro Kohno for providing the line ratios of galaxies used in Figure 13. H.-A.P. thanks Dr. Akihiko Hirota for helping with the improvement in English. We thank Dr. Eva Schinnerer for providing images in FITS format, although not ultimately used in this manuscript. The Submillimeter Array is a joint project between the Smithsonian Astrophysical Observatory and the Academia

Sinica Institute of Astronomy and Astrophysics and is funded by the Smithsonian Institution and the Academia Sinica. The Australia Telescope is funded by the Commonwealth of Australia for operation as a National Facility managed by CSIRO. We thank the staff of the ATNF for their efforts in development and characterization of the ATCA 3 mm system. Gemini Observatory is operated by the Association of Universities for Research in Astronomy, Inc., under a cooperative agreement with the NSF on behalf of the Gemini partnership: the National Science Foundation (United States), the Science and Technology Facilities Council (United Kingdom), the National Research Council (Canada), CONICYT (Chile), the Australian Research Council (Australia), Ministério da Ciência, Tecnologia e Inovação (Brazil), and Ministerio de Ciencia, Tecnología e Innovación Productiva (Argentina). This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

*Facilities:* ATCA, SMA, Gemini:South

## REFERENCES

- Aalto, S., Booth, R. S., Black, J. H., & Johansson, L. E. B. 1995, *A&A*, **300**, 369
- Allard, E. L., Peletier, R. F., & Knapen, J. H. 2005, *ApJ*, **633**, 25
- Athanassoula, E. 1992, *MNRAS*, **259**, 345
- Barnes, D. G., Staveley-Smith, L., de Blok, W. J. G., et al. 2001, *MNRAS*, **322**, 486
- Bigiel, F., Leroy, A., Walter, F., et al. 2008, *AJ*, **136**, 2846
- Böker, T., Falcón-Barroso, J., Schinnerer, E., Knapen, J. H., & Ryder, S. 2008, *AJ*, **135**, 479
- Brandl, B. R., Martín-Hernández, N. L., Schaerer, D., Rosenberg, M., & van der Werf, P. P. 2012, *A&A*, **543**, 61
- Buta, R., & Combes, F. 1996, *FCPh*, **17**, 95
- Claussen, M. J., & Sahai, R. 1992, *AJ*, **103**, 1134
- Combes, F., & Gerin, M. 1985, *A&A*, **150**, 327
- Condon, J. J. 1992, *ARA&A*, **30**, 575
- Condon, J. J., Cotton, W. D., & Broderick, J. J. 2002, *AJ*, **124**, 675
- Dahlem, M. 2005, *A&A*, **429**, 5
- Davies, R. I., Sugai, H., & Ward, M. J. 1998, *MNRAS*, **300**, 388
- Downes, D., Reynaud, D., Solomon, P. M., & Radford, S. J. E. 1993, *ApJ*, **414**, 13
- Downes, D., Reynaud, D., Solomon, P. M., & Radford, S. J. E. 1996, *ApJ*, **461**, 186
- Durret, F., & Bergeron, J. 1987, *A&A*, **173**, 219
- Ellison, S. L., Nair, P., Patton, D. R., et al. 2011, *MNRAS*, **416**, 2182
- Elmegreen, B. G. 1994, *ApJL*, **425**, L73
- Eskridge, P. B., Frogel, J. A., Pogge, R. W., et al. 2000, *AJ*, **119**, 536
- Feinstein, C., Mendez, M., Vega, I., & Forte, J. C. 1990, *A&A*, **239**, 90
- Forbes, D. A., Kotilainen, J. K., & Moorwood, A. F. M. 1994a, *ApJ*, **433**, 13
- Forbes, D. A., Norris, R. P., Williger, G. M., & Smith, R. C. 1994b, *AJ*, **107**, 984
- Gao, Y., & Solomon, P. M. 2004, *ApJ*, **606**, 271
- Goldreich, P., & Kwan, J. 1974, *ApJ*, **18**, 441
- Ho, P. T. P., Moran, J. M., & Lo, K. Y. 2004, *ApJ*, **616**, 1
- Hook, I., Jørgensen, I., Allington-Smith, J. R., et al. 2004, *PASP*, **116**, 425
- Hsieh, P. Y., Matsushita, S., Lim, J., Kohno, K., & Sawada-Satoh, S. 2008, *ApJ*, **683**, 70
- Hsieh, P. Y., Matsushita, S., Liu, G., et al. 2011, *ApJ*, **736**, 129
- Hüttemeister, S., Aalto, S., Das, M., & Wall, W. F. 2000, *A&A*, **363**, 93
- Irvine, W. M., Goldsmith, P. F., & Hjalmarsen, A. 1987, in *Interstellar Process*, ed. D. J. Hollenbach & H. A. Thronson (Dordrecht: Reidel), 585
- Jogee, S., Scoville, N., & Kenney, J. D. P. 2005, *ApJ*, **630**, 837
- Kenney, J. D. P., Wilson, C. D., Scoville, N. Z., Devereux, N. A., & Young, J. S. 1992, *ApJ*, **395**, 79
- Kennicutt, R. C., Jr. 1998, *ARA&A*, **36**, 189
- Kennicutt, R. C., Jr., Armus, L., Bendo, G., et al. 2003, *PASP*, **115**, 928
- Knapen, J. H. 2005, *A&A*, **429**, 141
- Knapen, J. H., Sharp, R. G., Ryder, S. D., et al. 2010, *MNRAS*, **408**, 797
- Kohno, K. 2005, in *AIP Conf. Proc.* 783, *The Evolution of Starbursts: The 331st Wilhelm and Else Heraeus Seminar* (Melville, NY: AIP), 203
- Kohno, K., Kawabe, R., & Vila-Vilaró, B. 1999, *ApJ*, **511**, 157
- Kohno, K., Matsushita, S., Vila-Vilaró, B., et al. 2001, in *ASP Conf. Ser.* 249, *The Central Kiloparsec of Starbursts and AGN: The La Palma Connection*, ed. J. H. Knapen, J. E. Beckman, I. Shlosman, & T. J. Mahoney (San Francisco, CA: ASP), 672
- Kohno, K., Nakanishi, K., & Imanishi, M. 2007, in *ASP Conf. Ser.* 373, *The Central Engine of Active Galactic Nuclei*, ed. L. C. Ho & J.-M. Wang (San Francisco, CA: ASP), 647
- Kohno, K., Nakanishi, K., Tosaki, T., et al. 2008, *Ap&SS*, **313**, 279
- Krips, M., Neri, R., García-Burillo, S., et al. 2007, *A&A*, **468**, 63
- Laurikainen, E., Salo, H., Buta, R., & Knapen, J. H. 2007, *MNRAS*, **381**, 401
- Leon, S., Jeyakumar, S., Pérez-Ramírez, D., et al. 2008, *A&A*, **491**, 703
- Malhotra, S. 1994, *ApJ*, **433**, 687
- Matsushita, S., Kawabe, R., Kohno, K., Tosaki, T., & Vila-Vilaro, B. 2010, *PASJ*, **62**, 409
- Matsushita, S., Kohno, K., Vila-Vilaro, B., Tosaki, T., & Kawabe, R. 1998, *ApJ*, **495**, 267
- Mazzuca, L. M., Knapen, J. H., Veilleux, S., & Regan, M. W. 2008, *ApJS*, **174**, 337
- Meier, D. S., & Turner, J. L. 2001, *ApJ*, **551**, 687
- Moorwood, A. F. M., & Oliva, E. 1988, *A&A*, **203**, 278
- Papadopoulos, P. P. 2010, *ApJ*, **720**, 226
- Papadopoulos, P. P., Thi, W.-F., Miniati, F., & Viti, S. 2011, *MNRAS*, **414**, 1705
- Regan, M. W., & Teuben, P. J. 2004, *ApJ*, **600**, 595
- Ryder, S. D., Illingworth, S. M., Sharp, R. G., & Farage, C. L. 2010, *PASA*, **27**, 56
- Ryder, S. D., Knapen, J. H., & Takamiya, M. 2001, *MNRAS*, **323**, 663
- Saikia, D. J., Phookun, B., Pedlar, A., & Kohno, K. 2002, *A&A*, **383**, 98
- Sakamoto, K., Okumura, S. K., Ishizuki, S., & Scoville, N. Z. 1999, *ApJS*, **124**, 403
- Sanders, D. B., Mazzarella, J. M., Kim, D.-C., Surace, J. A., & Soifer, B. T. 2003, *AJ*, **126**, 1607
- Schinnerer, E., Eckart, A., Quirrenbach, A., et al. 1997, *ApJ*, **488**, 174
- Schinnerer, E., Maciejewski, W., Scoville, N., & Moustakas, L. A. 2002, *ApJ*, **575**, 826
- Schmitt, H. R., Calzetti, D., Armus, L., et al. 2006, *ApJS*, **164**, 52
- Scoville, N. Z., Carlstrom, J. E., Chandler, C. J., et al. 1993, *PASP*, **105**, 1482
- Scoville, N. Z., & Solomon, P. M. 1974, *ApJL*, **187**, L67
- Sheth, K., Vogel, S. N., Regan, M. W., et al. 2002, *AJ*, **124**, 2581
- Shibatsuka, T., Matsushita, S., Kohno, K., & Kawabe, R. 2003, *PASJ*, **55**, 87
- Smith, D. A., Herter, T., Haynes, M. P., & Neff, S. G. 1999, *ApJ*, **510**, 669
- Sofue, Y. 1996, *ApJ*, **458**, 120
- Sofue, Y. 1997, *PASJ*, **49**, 17
- Solomon, P. M., Downes, D., & Radford, S. J. E. 1992, *ApJL*, **387**, L55
- Solomon, P. M., Rivolo, A. R., Barrett, J., & Yahil, A. 1987, *ApJ*, **319**, 730
- Solomon, P. M., Scoville, N. Z., & Sanders, D. B. 1979, *ApJL*, **232**, L89
- Stark, A. A., & Brand, J. 1989, *ApJ*, **339**, 763
- Taniguchi, Y., & Ohya, Y. 1998, *ApJL*, **507**, L121
- Toomre, A. 1964, *ApJ*, **139**, 1217
- Watson, W. D., Anicich, V. G., & Huntress, W. T. 1976, *ApJ*, **205**, 165