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HST-COS OBSERVATIONS OF HYDROGEN, HELIUM, CARBON AND NITROGEN EMISSION FROM THE SN 1987A REVERSE SHOCK*

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ABSTRACT

We present the most sensitive ultraviolet observations of Supernova 1987A to date. Imaging spectroscopy from the *Hubble Space Telescope*-Cosmic Origins Spectrograph shows many narrow ($\Delta v \sim 300$ km s⁻¹) emission lines from the circumstellar ring, broad ($\Delta v \sim 10 - 20 \times 10^3$ km s⁻¹) emission lines from the reverse shock, and ultraviolet continuum emission. The high signal-to-noise (> 40 per resolution element) broad Ly α emission is excited by soft X-ray and EUV heating of mostly neutral gas in the circumstellar ring and outer supernova debris. The ultraviolet continuum at $\lambda > 1350$ Å can be explained by H I 2-photon ($2s^2S_{1/2} - 1s^2S_{1/2}$) emission from the same region. We confirm our earlier, tentative detection of N V $\lambda 1240$ emission from the reverse shock and present the first detections of broad He II $\lambda 1640$, C IV $\lambda 1550$, and N IV] $\lambda 1486$ emission lines from the reverse shock. The helium abundance in the high-velocity material is He/H = 0.14 ± 0.06 . The N V/H α line ratio requires partial ion-electron equilibration ($T_e/T_p \approx 0.14 - 0.35$). We find that the N/C abundance ratio in the gas crossing the reverse shock is significantly higher than that in the circumstellar ring, a result that may be attributed to chemical stratification in the outer envelope of the supernova progenitor. The N/C abundance ratio may have been stratified prior to the ring expulsion, or this result may indicate continued CNO processing in the progenitor subsequent to the expulsion of the circumstellar ring.

Subject headings: supernovae: individual (SN 1987A) — shock waves — circumstellar matter

1. INTRODUCTION

Borkowski, Blondin, & McCray (1997) predicted that the spectrum of SN 1987A should display very broad ($\Delta v \sim \pm 12,000$ km s⁻¹) emission lines of Ly α , H α , N V $\lambda 1240$, and He II $\lambda 1640$, produced where the freely expanding supernova debris crosses a reverse shock located inside the equatorial circumstellar ring. In September 1997, using the *Hubble Space Telescope*-Space Telescope Imaging Spectrograph (*HST*-STIS), Sonneborn et al. (1998) detected broad Ly α emission, and Michael et al. (1998) showed how observations of this emission can be used to map the shape of the reverse shock and the flux of H I atoms crossing it. Michael et al (2003) and Heng et al. (2006) analyzed subsequent (February 1999 – October 2002) STIS observations of both Ly α and H α to map the increasing flux of H I atoms across the reverse shock.

The evolution of the broad H α emission can also be tracked with ground-based telescopes, notably in February 2005 with the Magellan telescope (Smith et al. 2005) and from December 2000 to January 2009 with the Very Large Telescope (Fransson et al 2011). Smith et al. pointed out that the H α emission from the reverse shock could be suppressed due to preionization of hydrogen in the supernova debris by soft X-rays and EUV radiation from the rapidly brightening shock interaction with the inner circumstellar ring. Extrapolating the X-ray light

curve, they predicted that this preionization would cause the broad H α emission to vanish by 2012. However, since 2005, the X-ray light curve has leveled off (Park et al. 2011), so the preionization effect should be less than that predicted by Smith et al.

France et al. (2010) analyzed the most recent (31 January 2010) STIS observations of SN 1987A. They pointed out that the observed ratio of Ly α /H α photon fluxes was $\gtrsim 30$ at large, negative velocities, much greater than the value ≈ 5 that would be expected for hydrogen atoms excited as they cross the reverse shock. Moreover, the spatially resolved Ly α line profile differed dramatically from that of H α . France et al. proposed that the broad Ly α emission is dominated by Ly α emission from the nearly stationary equatorial ring that has been resonantly scattered by hydrogen atoms in the expanding supernova debris. France et al. also noticed a faint glow at wavelengths ranging from about 1260 – 1290 Å, which they attributed to the N V $\lambda 1240$ emission predicted by Borkowski et al. They suggested that a critical test of that hypothesis would be detection of broad emission from C IV $\lambda 1550$.

Here we describe far-ultraviolet (1140 – 1780 Å) spectra of SN 1987A obtained with the Cosmic Origins Spectrograph (COS) that was installed on *HST* in May 2009. With COS, we have measured the profiles of Ly α and N V $\lambda 1240$ with signal-to-noise ratios (S/N) far superior to the previous STIS spectra. We have also detected the broad He II $\lambda 1640$ emission predicted by Borkowski et al. (1997) and the broad C IV $\lambda 1550$ emission predicted by France et al (2010). We find that while the He/H abundance ratio in the outer ejecta is consistent with that

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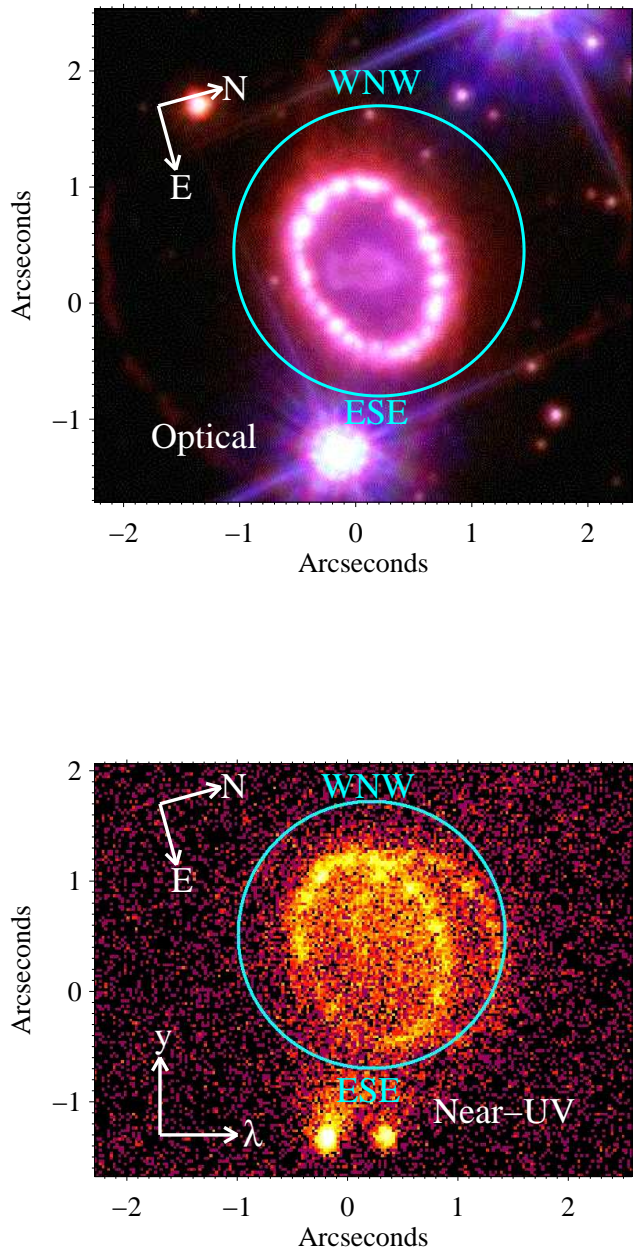


FIG. 1.— (top) *HST*-ACS *BVR* image with a representative overlay of the 2.5'' diameter COS aperture (cyan). (bottom) *HST*-COS near-UV pointing verification image for the G130M (2011 February 11) visit. Due to bright limits, this observation was made using the MIRRORB imaging configuration, which creates the image doubling. Because COS is a slitless spectrograph, light from outside the nominal 2.5'' can enter the system. Star 3, separated from the center of the circumstellar ring by $\approx 1.7''$ is seen at the bottom. In order to match the spectral image presented in Figure 2, we display the coordinates of this image with east to the right of north, opposite from the standard convention. The spectroscopic dispersion axis is labeled “ λ ” and the cross-dispersion axis is labeled “ y ”.

of the circumstellar ring, the N/C ratio is enhanced by more than a factor of two over the ring abundances.

2. *HST*-COS OBSERVATIONS, IMAGING SPECTROSCOPY, AND DATA REDUCTION

SN 1987A was observed with the medium resolution far-UV modes of *HST*-COS (G130M and G160M) on 2011 February 11 and March 14 for a total of 7 spacecraft orbits (18555 s; Table 1) as part of the Supernova 1987A INTensive Study (SAINTS - GO12241; PI - R. Kirshner). A description of the COS instrument and on-orbit performance characteristics can be found in Osterman et al. (2011). All observations were approximately centered on the SN 1987A circumstellar ring (R.A. = $05^{\text{h}} 35^{\text{m}} 28.07^{\text{s}}$, Dec. = $-69^{\circ} 16' 10.8''$; J2000) and COS performed an offset imaging target acquisition from a reference star $\approx 8''$ away. A NUV image was obtained with the MIRRORB imaging mode after the slew to the science pointing in order to verify that the circumstellar ring was in the primary science aperture (PSA). Figure 1 (top) shows the COS aperture location on an optical image of SN 1987A. The MIRRORB configuration introduces optical distortions into the image, but a first-order analysis indicates that the entire circumstellar ring was inside the 2.5'' diameter PSA (Figure 1, bottom).

The COS position angle was chosen to maximize the spatial separation on the COS detector between Star 3 and the supernova emission we aim to study. The G130M observations were made at a position angle of $\approx -75^{\circ}$, while the G160M observations were acquired with a position angle of $\approx -45^{\circ}$. Light from objects outside the nominal 1.25'' COS aperture radius can be recorded with the science spectrum in crowded fields. Star 3 is a far-UV-bright Be star (Gilmozzi et al. 1987; Walborn et al. 1993) separated from the center of the SN 1987A ring by $\approx 1.7''$. Stellar contamination could compromise supernova data quality if the two objects overlapped on the spectroscopic detector. Thus, we used the imaging spectroscopic capability of COS to keep SN 1987A centered while placing Star 3 at the bottom of the microchannel plate (MCP) detector.

We complemented this observing strategy with post-processing techniques and custom spectral extractions to maximize the spatial resolution of the instrument in the cross-dispersion direction. Due to long-term exposure of the COS MCP to geocoronal Ly α and hot-star spectra, the number of electrons generated by each incident photon at a given detector location has been decreasing. This is manifest as lower pulse heights observed by the detector electronics². The exact form of the pulse-height evolution is a complicated function of the dispersion and cross-dispersion arrival location of the incident photon, but to first order we can approximate this as a pulse-height dependent mislocation of the photon y -location. The approximate form of this mislocation is $-0.47 y$ pixels per pulse height bin drop. We have reprocessed the COS observations of SN 1987A, including a correction for the cross-dispersion misplacement, producing more accurate photon arrival positions. This additional data reduction step improves the far-UV angular resolution of the instrument by 20 – 30% across the bandpass.

Star 3 can be seen at the bottom of the two-dimensional spectrogram shown in Figure 2. Star 3 is well isolated at the bottom of the detector. The angular resolution in the reprocessed two-dimensional data

² We refer the reader to the cycle 19 COS Instrument Handbook for more details: http://www.stsci.edu/hst/cos/documents/handbooks/current/cos_cover.htm

is $\approx 0.8''$. This imaging capability allows us to make custom spectral extractions at three spatially resolved locations: the “WNW” region at the top of the detector (approximately centered on the hotspots at: R.A. = $05^{\text{h}} 35^{\text{m}} 27.98^{\text{s}}$, Dec. = $-69^{\circ} 16' 10.5''$; J2000), the “ESE” region towards Star 3 (approximately the telescope pointing coordinates), and a separate extraction of Star 3 itself. Star 3 contributes negligible flux to the one-dimensional WNW spectrum and less than 15% of the flux in the ESE extraction. The spatially resolved spectra will be analyzed in a future work, and we focus on the combined WNW + ESE spectra here.

The custom data extractions were then reprocessed with the COS calibration pipeline, CALCOS v2.13.6, and combined with the custom IDL coaddition procedure described by Danforth et al. (2010) and Shull et al. (2010). The coaddition routine interpolates all detector segments and grating settings onto a common wavelength grid, and makes a correction for the detector QE-enhancement grid. No correction for the detector hex pattern is performed. Data were obtained in four central wavelength settings in each far-UV grating mode ($\lambda 1291$, 1300, 1309, and 1318 with G130M and $\lambda 1577$, 1589, 1600, and 1611 with G160M) at the default focal-plane split position. The instrumental configurations are summarized in Table 1. Observations at multiple wavelength settings provide continuous spectral coverage over the 1136 – 1782 Å bandpass and minimize the residual fixed pattern noise from the detector grid wires and the MCP pores. The point source resolving power of the medium resolution COS far-UV modes is $R \equiv \Delta\lambda/\lambda \approx 18,000$ ($\Delta v = 17 \text{ km s}^{-1}$); however, the filled-aperture resolving power is $R \sim 1500$ ($\Delta v \sim 200 \text{ km s}^{-1}$; France et al. 2009). Multiple point sources within the aperture (e.g., the multiple hotspots along the SN 1987A circumstellar ring) produce spectral resolution that is similar to the extended source response. The point source flux calibration of COS is accurate to better than $\approx 3\%$; however, the extended source and custom processing limit the absolute flux cal-

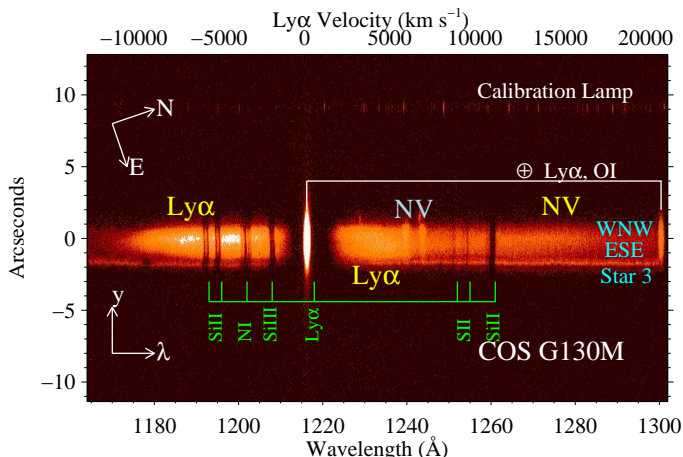


FIG. 2.— Two-dimensional spectrogram of the G130M segment B data. Reverse shock emission is labeled in yellow, hotspot emission is labeled in blue, and interstellar absorption features are marked in green.

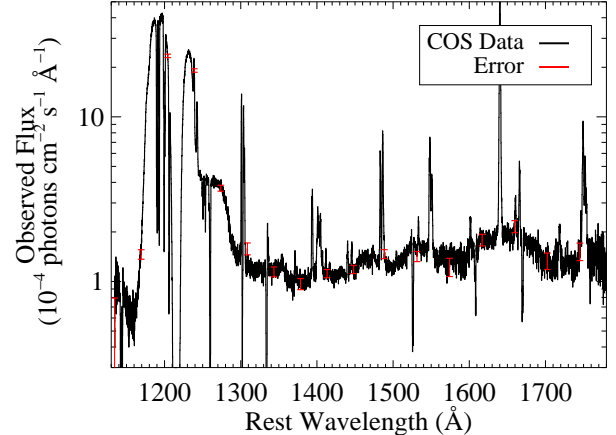


FIG. 3.— Full far-UV spectrum of the combined WNW and ESE regions. Broad emission from the SN 1987A reverse shock, narrow emission lines attributable to circumstellar hotspots, as well as underlying continuum emission are observed. Representative $1\text{-}\sigma$ error bars (a combination of photon statistics and flux calibration uncertainties) are shown in red.

ibration of the SN 1987A far-UV spectroscopic data to $\sim 10\%$.

3. DATA ANALYSIS

3.1. Emission Spectrum and Interstellar Corrections

The extracted emission spectrum of SN 1987A is shown in Figure 3. The spectra presented here are shown in photon units as these are the natural units for line flux comparisons. There is a wealth of narrow and broad emission and absorption features in the spectrum as well as a strong underlying continuum. The “narrow” (FWHM $\approx 200 - 300 \text{ km s}^{-1}$) resolved emission features are attributed to the circumstellar ring. The circumstellar ring spectrum is dominated by the hotspots which originate at shock interfaces where the forward blast wave encounters regions of high density in the circumstellar ring (Lawrence et al. 2000). A discussion of the ring-ejecta interaction is given by Larsson et al. (2011). Most of the circumstellar ring emission lines have been observed in previous *HST*-STIS (e.g., Pun et al. 2002) and *IUE* (Lundqvist & Fransson 1996) observations, including N V $\lambda 1239$, 1243, C II $\lambda 1334$, 1335, Si IV $\lambda 1394$, 1403, the O IV] $\lambda 1400$ multiplet, N IV] $\lambda 1483$, 1486, C IV $\lambda 1548$, 1550, [Ne IV] $\lambda 1602$, He II $\lambda 1640$, O III] $\lambda 1661$, 1666, and the N III] $\lambda 1750$ multiplet. Our deep COS observations enable first detections of O V $\lambda 1371$ and Si II $\lambda 1526$, 1533, although the ground state transition of Si II is blocked by the interstellar medium. We also tentatively detect the coronal [Si VIII] doublet $\lambda 1441$, 1446. These features will be discussed in detail in a future work, though we will make use of the circumstellar ring line fluxes where relevant.

The broad emission features observed in the spectrum of SN 1987A arise as atoms and ions cross the reverse shock front (Michael et al. 1998), located just interior to the circumstellar ring. H I Ly α (and the corresponding H α) emission from the reverse shock has been well studied (Michael et al. 2003; Smith et al. 2005; Heng et al. 2006; France et al. 2010), and we display the

TABLE 1
SN 1987A COS OBSERVING LOG.

Date	COS Mode	Central Wavelengths	FP-POS	T_{exp} (s)
2011 February 11	G130M	1291,1300,1309,1218	3	7884
2011 March 14	G160M	1577,1589,1600,1611	3	10671

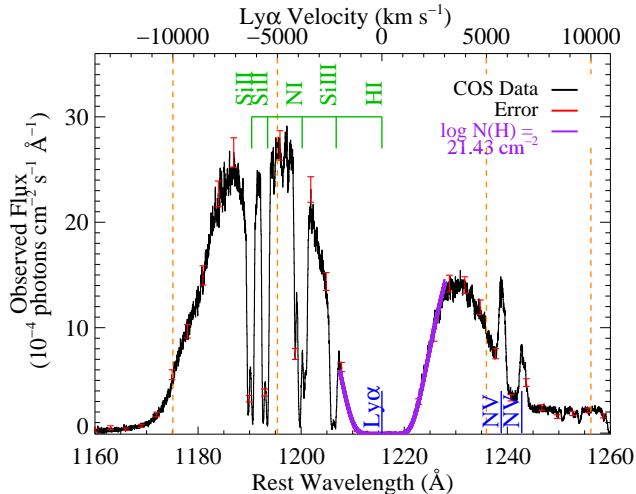


FIG. 4.— RS Ly α and hotspot N V in the WNW extraction. Hotspot emission lines are marked in blue and interstellar absorption components are labeled in green. The purple line shows a fit to the Ly α line core for $\log_{10}(N(\text{H})) = 21.43 (\pm 0.02) \text{ cm}^{-2}$. Dashed orange lines are plotted as velocity references. Representative $1-\sigma$ error bars (a combination of photon statistics and flux calibration uncertainties) are shown in red.

broad ($\Delta v \sim -13,000 - +7,000 \text{ km s}^{-1}$) Ly α emission seen in the COS observations in Figures 4 and 5. The second broad feature to the red of Ly α in Figure 3 is consistent with the putative N V reverse shock emission described by France et al. (2010) and has a velocity distribution and total flux that are approximately consistent with those predicted by Borkowski et al. (1997); see also §4.3.1. In light of these results and the large relative offset from the observed Ly α profile, we claim that reverse shock N V emission is unambiguously detected in our data (Figure 5). While at earlier times, low-ionization UV line emission from SN 1987A has been attributed to the debris in the core (Jerkstrand et al. 2011), the core is enshrouded in dust (Matsuura et al. 2011) and is mostly opaque to UV radiation. Therefore, the reverse shock is the most likely formation site for high-velocity UV line photons observed towards SN 1987A. Following similar line-identification arguments as those for N V, we also detect reverse shock emission from N IV] $\lambda 1486$, C IV $\lambda 1550$, and He II $\lambda 1640$ for the first time. These features are marked in Figure 6.

Figure 4 also shows an example of the broad and narrow interstellar absorption features imposed on the spectrum of SN 1987A. This figure shows Galactic and Magellanic Cloud absorption from Si II $\lambda 1190$, 1193, the N I $\lambda 1200$ multiplet, Si III $\lambda 1206$, and H I Ly α . Other interstellar absorbers include S II $\lambda 1250$, 1253, Si II $\lambda 1260$, C II $\lambda 1334$, Si II $\lambda 1526$, Fe II $\lambda 1608$, and Al II $\lambda 1670$. We fit the interstellar neutral hydrogen absorption in the WNW spectrum as this provides the best measure of the foreground absorption, finding

$\log_{10}(N(\text{H})) = 21.43 \pm 0.02 \text{ cm}^{-2}$. For subsequent analysis of the Ly α and N V emission, we divide the data by a model of this emission (shown in purple in Figure 4). This procedure renders unusable the inner $\pm 6 \text{ \AA}$ of the Ly α line profile, which we remove from the data.

There are several estimates of the interstellar reddening towards the region of the LMC in which SN 1987A resides (Walker & Suntzeff 1990; Fitzpatrick & Walborn 1990; Scuderi et al. 1996; Michael et al. 2003). We adopt a Milky Way ($R_V = 3.1$) extinction curve (Cardelli et al. 1989) with $E(B - V) = 0.19$ as this curve approximates the average extinction correction favored by the studies cited above. This curve also closely approximates the far-UV properties of the average LMC curve presented by Gordon et al. (2003) for $A_V = 0.6$. Assuming $E(B - V) = 0.17$ (Michael et al. 2003), would lower the fluxes derived in subsequent sections by $\sim 15\%$. Scuderi et al. (1996) present a comprehensive study of the extinction towards Star 2, and the use of this curve would result in fluxes $\sim 20 - 30\%$ larger than those derived below. None of the possible extinction curves significantly alters the shape of the continuum (§4.1.1) or the far-UV line ratios (§4.3.3).

The bright continuum ($F_\lambda(\text{observed}) \sim 1 \times 10^{-4} \text{ photons cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$) underlying the hotspot and reverse shock emission was suggested in coadded STIS G140L spectra from 2010, cospatial with the circumstellar ring. However, the significance of the detection was low. We propose that this continuum is mostly H I 2-photon emission (§4.1.1).

3.2. Binned Spectrum for Broad Line & Continuum Analysis

Narrow emission lines, as well as the narrow interstellar absorption components, complicate measurements of the continuum and broad emission from the reverse shock. In order to remove confusion from these narrow lines during the analysis of the broad spectral features, we created binned spectra of the WNW and ENE extractions. These binned spectra were centered on hand-chosen regions of the data at $1 - 5 \text{ \AA}$ intervals and were free of narrow features. The flux and error of these binned data points were taken to be the average flux and standard deviation of a 0.8 \AA region centered on the selected wavelength. These binned data are shown as diamonds overplotted on the spectra in Figures 5 – 8.

We quantify the total reverse shock flux by integrating the binned spectra over wavelength intervals corresponding to the velocity ranges of interest for a given feature. The velocity ranges were chosen to cover the maximum extent of a given line without significantly overlapping with other reverse shock emission features. The inner approximately 3000 km s^{-1} were not included to avoid strong hotspot emission lines located at $\approx 0 \text{ km s}^{-1}$. In most cases the hot spot lines are doublets, therefore a velocity interval larger than the nominal 300 km s^{-1} nar-

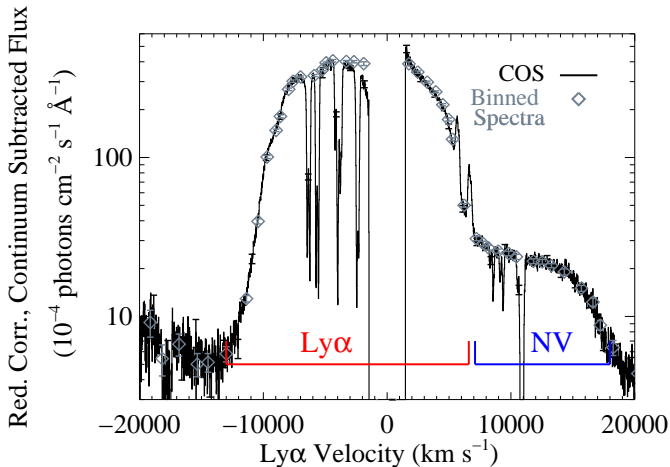


FIG. 5.— Velocity profile of the ISM corrected (dust and neutral hydrogen), continuum subtracted RS Ly α and N V in the combined WNW + ESE spectrum. The binned spectrum described in §3.2 are shown overlaid in gray. The Ly α emission extends from -13000 – $+7000$ km s $^{-1}$ and emission to the red of Ly α is attributed to N V.

row line-width was excluded. We were unable to measure two reverse shock components due to velocity blending. The blue component of the N V line is lost under the Ly α line, and the weak red wing of the N IV] emission is overwhelmed by the stronger blue C IV emission. The velocity and wavelength intervals are given in Table 2 with the summed line fluxes.

4. RESULTS

4.1. Hydrogen: 2-photon Emission and Broad Ly α

4.1.1. H I Continuum

Figure 3 shows that the emission and absorption lines are superposed upon a far-UV continuum with $F_{\lambda} > 10^{-4}$ photons cm $^{-2}$ s $^{-1}$ Å $^{-1}$. This continuum was marginally detected in previous STIS G140L spectra, with the emission being concentrated in the ring plane. It is detected in our COS observations at high significance, ($Flux / Error$) $\gtrsim 8$ per spectral resolution element from 1360 – 1540 Å.

We can rule out the possibility that this emission is uncorrected scattered light from Star 3. As described in Section 2, care was taken to isolate the emission from Star 3 at the bottom of the MCP detector (Figure 2). At the observed continuum levels, the scattered light from Star 3 would have to be over twice as bright outside of the stellar extraction region as within it, and relatively constant with angular separation from the stellar spectrum. Additionally, the observed continuum of Star 3 has a different spectral shape than the supernova continuum and there is no evidence for stellar features in the supernova spectrum (such as photospheric and wind absorption lines of C III $\lambda 1175$ and Si IV $\lambda 1394$, 1403; Pellerin et al. 2002). Therefore, we consider it unlikely that Star 3 contributes more than 10% to the supernova continuum at $\lambda > 1350$ Å.

We propose that this continuum is predominantly H I 2-photon ($2s^2S_{1/2} - 1s^2S_{1/2}$) emission. We fit the $1350 \leq \lambda \leq 1720$ Å spectrum with the analytic formula

of Nussbaumer & Schmutz (1984),

$$J_{\lambda} = \frac{1}{4\pi} \frac{hc}{\lambda^3} A(\lambda) N_{2s^2S} \quad (1)$$

where J_{λ} is the emissivity and N_{2s^2S} is the H I column density in the $2s^2S$ state. This function rises sharply from 1216 Å to a maximum at $\lambda \sim 1420$ Å, and then declines slowly to the red. The 2-photon spectral model is shown in orange in Figure 6. Based on our near-UV images, we estimate that the circumstellar ring fills approximately $30 \pm 20\%$ of the inner $2''$ of the COS aperture where the transmission is > 0.6 . Assuming this filling fraction, the continuum level determines the column density, and we find $\log_{10}(N_{2s^2S}(\text{H})) = 21.34^{+0.50}_{-0.26}$ cm $^{-2}$.

The total integrated (1216 Å – 6 μm) H I 2-photon flux is $9.1 \pm 0.6 \times 10^{-1}$ photons cm $^{-2}$ s $^{-1}$, and the observed ratio of Ly α line emission to 2-photon is $F(\text{Ly}\alpha)/F(2s) = 1.96 \pm 0.23$. Since this ratio is consistent with the $2p/2s$ ratio of 2.1 (observed as $F(\text{Ly}\alpha)/F(2s)$) expected for recombination at $T \approx 10^4$ K (Spitzer 1978) and the $(1s \rightarrow 2p)/(1s \rightarrow 2s)$ ratio of 2.05 – 2.10 expected for excitation by thermal electrons (Callaway 1988), we infer that it is likely that the Ly α and the 2-photon emission come from the same source.

An important constraint on the origin of the far-UV continuum is the density of the emitting region. H I 2-photon emission can be suppressed by a factor $[1 + n_p/n_{cr}]^{-1}$ due to collisional depopulation of the $2s$ state, where the critical density is given by $n_{cr} = A_{2\gamma}/C_{sp}$ and n_p is the density of protons. The $(2s \rightarrow 2p)$ collision rate coefficient is $C_{sp} = 5.31 \times 10^{-4}$ cm 3 s $^{-1}$ (Seaton 1955; see also Dennison et al. 2005 for a discussion of $2s$ and $2p$ level populations in H II regions) and $A_{2\gamma} \sim 8.23$ s $^{-1}$ for the 2-photon decay (Klarsfeld 1969), therefore $n_{cr} = 1.5 \times 10^4$ cm $^{-3}$.

Based on the observed 2-photon flux, we can predict the total amount of associated H α if this emission is predominantly recombination. In Case B recombination, the rate of population of the $2s$ state is 3.3 times the rate of emission of H α (Osterbrock & Ferland 2006). The expected H α in this scenario is $\approx 2.8 \times 10^{-1}$ photons cm $^{-2}$ s $^{-1}$. We have made an estimate of the total H α flux from the circumstellar ring by analyzing STIS G750M spectra obtained in 2009 October (*HST* exposure IDs OB7I200A0 – OB7I200D0), making a correction for differential reddening between H α and the 2-photon emission. The total H α emission from the ring is $\approx 2.5 \times 10^{-1}$ photons cm $^{-2}$ s $^{-1}$, very close to the value expected if the 2-photon emission was produced by recombination. However, Grönningsson et al. (2008b) find that the H α flux from the ring is dominated by high-density ($> 4 \times 10^6$ cm $^{-3}$) shocked material around 5700 days after the explosion (October 2002). 2-photon emission could be highly suppressed from such gas, which has $n_p/n_{cr} \gtrsim 300$. 2-photon emission from lower density ($\lesssim 5 \times 10^3$ cm $^{-3}$; Matilla et al. 2010) gas in the unshocked ring would not be suppressed, but such gas can only account for $\sim 20\%$ of the observed 2-photon emission.

Because we observe no comparable source of H α emission, we conclude that the majority of the 2-photon emission is produced not by recombination, but by thermal

($T \approx 10^4$ K) electron impact excitation of the $2s^2S$ state of neutral hydrogen atoms in the outer ejecta, near the reverse shock front. This low-density gas ($n_H \sim 100 \text{ cm}^{-3}$; Smith et al. 2005; Heng et al. 2006) is heated by X-rays emitted from the shocked gas near the hotspots. In particular, the energy deposition of the soft X-ray/EUV photons from the shocked hotspots will be concentrated in a layer near the reverse shock (Fransson et al. 2011). When the ionization fraction is $\gtrsim 3 \times 10^{-2}$, the majority of the X-ray energy heats the gas through Coulomb stopping of fast photoelectrons (Xu et al. 1992; Kozma & Fransson 1992). The primary coolants for this gas are $\text{Ly}\alpha$ and 2-photon emission.

Is it reasonable to expect this level of X-ray heating of gas near the reverse shock? We calculate the total X-ray flux from the circumstellar ring using the two-component model spectrum of Zhekov et al. (2006), scaled to the total 0.5 – 2.0 keV luminosity observed by *Chandra* near day 8000 ($L(0.5 - 2.0 \text{ keV}) \approx 1.5 \times 10^{36} \text{ erg s}^{-1}$; Racusin et al. 2009). The total 2-photon luminosity is $L(2s) \approx 2.2 \times 10^{36} \text{ erg s}^{-1}$, therefore the 0.5 – 2.0 keV X-ray flux is insufficient to power the far-UV continuum. However, the total shock luminosity is most likely dominated by emission in the soft X-ray/EUV band (0.01 – 0.5 keV) that is attenuated by the neutral hydrogen in the interstellar media of the Milky Way and LMC (Fransson et al. 2011). The luminosity in this band inferred from the model by Zhekov et al. is $L(0.01 - 0.5 \text{ keV}) \approx 3 \times 10^{38} \text{ erg s}^{-1}$. Assuming that roughly half of this emission intersects the outer ejecta, we find that approximately 1.4% of the soft X-ray/EUV luminosity from the shocked ring must be reprocessed into H I 2-photon emission.

4.1.2. H I Line Emission

Figures 4 and 5 show the broad emission from H I $\text{Ly}\alpha$. This emission was first observed in *HST*-STIS spectra 10.25 years after the SN 1987A explosion (Sonneborn et al. 1998). Subsequent studies have explored the reverse shock $\text{Ly}\alpha$ and $\text{H}\alpha$ emission in more detail (Michael et al. 1998, 2003), including its brightening and deceleration (Heng et al. 2006; see also Smith et al. 2005 and Fransson et al. 2011 for discussion of the $\text{H}\alpha$ evolution). The combination of increased sensitivity and spectral resolution and low instrumental background of COS enable us to produce the highest-quality ultraviolet velocity profiles of the reverse shock emission to date. The velocity distribution of the neutral hydrogen emission extends from $-13000 - +8000 \text{ km s}^{-1}$ (Figure 5). The velocity maxima are much smaller than the initial observations by Sonneborn et al. (1998; $\pm 20000 \text{ km s}^{-1}$) and are consistent with the decrease in maximum projected velocity observed from 2004 to 2010 (France et al. 2010). Using the limited angular resolution of COS, we confirm that emission at negative velocities is concentrated towards the northern side of the circumstellar ring while positive velocities are concentrated on the southern side of the ring. The total integrated ($-12000 - +8000 \text{ km s}^{-1}$) $\text{Ly}\alpha$ flux inferred from Figure 5 is $1.78 \pm 0.18 \text{ photons cm}^{-2} \text{ s}^{-1}$.

Previous work noted that the $\text{Ly}\alpha/\text{H}\alpha$ ratios from the reverse shock exceed the 5:1 photon production ratio expected for a Balmer-dominated shock (Heng et al.

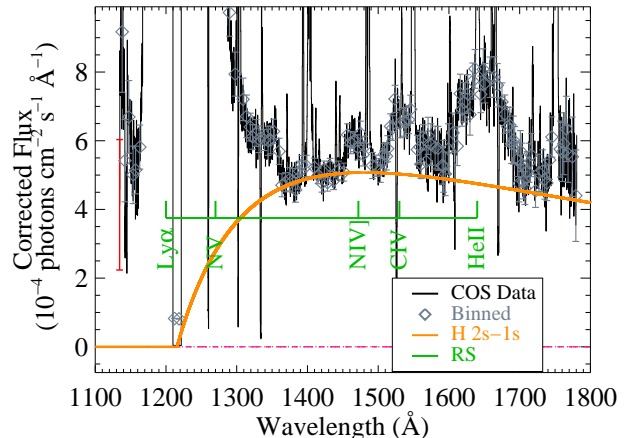


FIG. 6.— The combined WNW and ENE spectrum of SN 1987A (black) and the spectrum averaged over 0.8 \AA bins chosen to avoid narrow features from hotspot emission and interstellar absorption. These data have been corrected for interstellar neutral hydrogen absorption ($\log_{10}(N(\text{H})) = 21.43$) and interstellar dust extinction, assuming $E(B - V) = 0.19$ and $R_V = 3.1$. The broad spectral features are produced by a combination of hydrogen 2-photon emission ($2s \rightarrow 1s$) and ionic emission from the reverse shock. A theoretical hydrogen 2-photon spectrum is shown overplotted in orange. Reverse shock emission features are identified in green.

2006; Heng & McCray 2007). France et al. (2010) found $\text{Ly}\alpha/\text{H}\alpha$ ratios ≥ 30 from $-8000 - -2500 \text{ km s}^{-1}$ and ≥ 20 from $+3000 - +7000 \text{ km s}^{-1}$ for isolated cuts across the northern and southern sides of the reverse shock, respectively. They attributed the enhanced $\text{Ly}\alpha$ emission to a second source of $\text{Ly}\alpha$ photons. They argued that $\text{Ly}\alpha$ photons from the hotspots are resonantly scattered by onrushing hydrogen with a distribution of velocities spanning a width $\Delta v_H \sim 3000 - 9000 \text{ km s}^{-1}$ (unlike $\text{Ly}\alpha$, $\text{H}\alpha$ is not a resonance line and therefore hotspot $\text{H}\alpha$ photons pass through the debris freely).

While the $\text{Ly}\alpha$ enhancement at the largest negative velocities can likely be explained by this mechanism, our preceding discussion of the 2-photon continuum suggests that the majority of the lower velocity ($\pm 7000 \text{ km s}^{-1}$) $\text{Ly}\alpha$ excess is attributable to the X-ray heating of the outer supernova debris. As described above, $\text{Ly}\alpha$ is the primary coolant for mostly neutral hydrogen-rich gas excited by soft X-ray/EUV photoelectrons, and the agreement between the observed and theoretical $F(\text{Ly}\alpha)/F(2s)$ ratio therefore argues that thermal electron collisions dominate the production of $\text{Ly}\alpha$ photons in the outer ejecta near the reverse shock front. Resonant scattering will cause the newly produced $\text{Ly}\alpha$ photons to emerge preferentially in the outward direction, substantially favoring the blue-shifted wing as observed.

4.2. He II $\lambda 1640$ Emission

The broad emission feature that we attribute to He II $\lambda 1640$ from the reverse shock is easily visible above the $1600 - 1700 \text{ \AA}$ continuum in Figure 6. In Figure 7, we compare the He II velocity profile with the red side of the N V $\lambda 1240$ (§4.3.1) profile. As neutral or singly ionized helium crosses the reverse shock, collisions with thermal electrons (Laming et al. 1996) cause the helium to emit

TABLE 2
 CONTINUUM SUBTRACTED SN 1987A REVERSE SHOCK EMISSION

Species	$\Delta\lambda_{obs}$ (Å)	Δv^a (km s ⁻¹)	Line Flux (photons cm ⁻² s ⁻¹)	Flux Uncertainty (photons cm ⁻² s ⁻¹)
H I ^b	1169.6 – 1208.0	-12000 – -1500	1.1×10^0	5.9×10^{-3}
H I	1222.8 – 1247.1	1500 – 8000	4.5×10^{-1}	4.1×10^{-3}
N V ^c	1197.9 – 1233.9	-10000 – -1000	5.7×10^{-2}	6.2×10^{-3}
N V	1247.1 – 1285.5	2000 – 12000	8.0×10^{-2}	8.7×10^{-4}
N IV]	1428.3 – 1478.8	-12000 – -1500	1.8×10^{-3}	4.9×10^{-4}
N IV] ^d	1497.0 – 1523.8	1500 – 8000
C IV	1491.9 – 1539.4	-12000 – -1500	4.6×10^{-3}	5.4×10^{-4}
C IV	1558.9 – 1584.6	1500 – 8000	2.1×10^{-3}	4.8×10^{-4}
He II	1590.4 – 1633.7	-10000 – -1000	6.8×10^{-3}	7.8×10^{-4}
He II	1658.0 – 1705.0	2000 – 12000	9.6×10^{-3}	8.1×10^{-4}

^a $v = 0$ assumed to be the rest wavelength of the stronger lines of the N V and C IV doublets.

^b Corrected for interstellar neutral hydrogen absorption of $\log_{10}(N(H)) = 21.43$, however the inner ± 6 Å (± 1500 km s⁻¹) cannot be reconstructed. All reverse shock fluxes have been corrected for interstellar dust extinction, assuming $E(B - V) = 0.19$ and $R_V = 3.1$.

^c Blue N V flux is overwhelmed by reverse shock Ly α emission and interstellar H I absorption. We assume the N V Blue/Red ratio is the same as the He II reverse shock emission (≈ 0.7).

^d Red N IV] flux is obscured by the stronger blue emission from C IV.

the analog to H α . The velocity distribution of the ionized helium is expected to be quite different from that of neutral hydrogen because the He⁺ ions that have crossed the reverse shock will be deflected by turbulent magnetic fields in the shock isotropization zone (Michael et al. 1998). The broad He II displays a somewhat asymmetric line shape, spanning approximately -9000 – $+11000$ km s⁻¹.

Interpolating the integrated red and blue reverse shock He II $\lambda 1640$ photon fluxes given in Table 2 through the narrow emission lines³, we measure a total flux of 1.9×10^{-2} photons cm⁻² s⁻¹. Because this line is the hydrogenic analog to H α , the $F(1640)/F(H\alpha)$ ratio should give a direct measure of the relative abundance of helium (by number), assuming full hydrogen and helium ionization and that our line identification and continuum subtraction are correct. We compare our He II measurement with the H α flux from the day ~ 8000 VLT-UVES observations presented by Fransson et al. 2011, $F_{obs}(H\alpha) \approx 1.4 \times 10^{-13}$ erg cm⁻² s⁻¹. We convert this value into a total reverse shock H α by applying a factor of 1.86 slit correction, a factor of 1.45 reddening correction at H α (Gröningson et al. 2008a), and a factor of 24/22 to account for the relative time of observation after the explosion. The integrated H α photon flux in 2011 is ≈ 0.14 photons cm⁻² s⁻¹. The He/H abundance ratio in the reverse shock is 0.14 ± 0.06 , in agreement with the He/H abundance ratio derived by Mattila et al. (2010) for the circumstellar ring. The fact that the He II $\lambda 1640/H\alpha$ ratio reproduces the circumstellar ring He/H abundance ratio suggests that the neutral hydrogen pre-ionization predicted by Smith et al. (2005) is negligible 22 years after the explosion.

Borkowski et al. (1997) predicted the time evolution of the observable (accounting for attenuation due to interstellar reddening) reverse shock flux in several abundant ions, including He II. Their predictions only extend to 2007, but we can extrapolate their curves to 24 years after the explosion for comparison with the COS data.

³ The total emission from the reverse shock is interpolated through the central region lost to narrow-line emission. In the case of He II, $F_{HeII}^{RS} = (1 + (3000/22000)) \times (F_{HeII}^{RS,red} + F_{HeII}^{RS,blue})$.

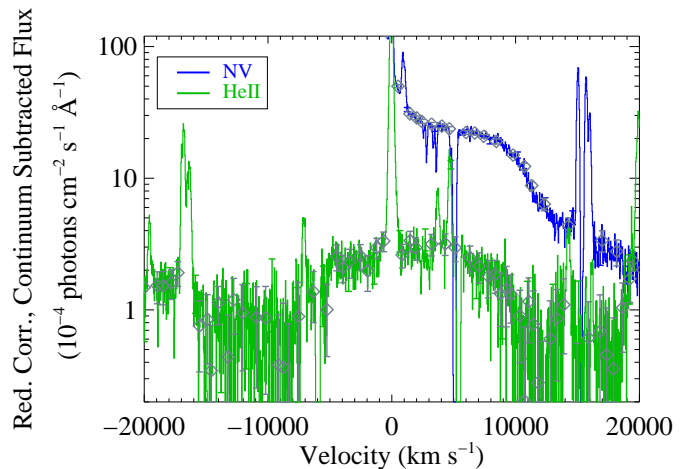


FIG. 7.— 2-photon continuum-subtracted N V and He II in the combined WNW + ESE spectra. The blue side of the N V distribution is lost under the RS Ly α emission. The He II velocity profile appears qualitatively similar to that of N V.

We estimate their prediction for the observed He II flux in 2011 to be $\sim 1 - 3 \times 10^{-3}$ photons cm⁻² s⁻¹ from Figure 4 of Borkowski et al. (1997). Correcting these values by a factor of 4 to account for interstellar reddening, their prediction for the He II flux is $0.4 - 1.2 \times 10^{-2}$ photons cm⁻² s⁻¹, only a factor of ~ 2 below the He II flux observed by COS. Differences between the predicted and observed He II fluxes are likely due to assumed electron temperatures that are lower than suggested by the observations (see the next subsection).

4.3. Carbon and Nitrogen Emission Lines

4.3.1. N V $\lambda 1240$

Prior to the first STIS observations of SN 1987A, Borkowski et al. (1997) predicted that strong reverse shock emission from Li-like N V $\lambda 1240$ would be detectable. This emission was not apparent, however, in the first deep far-UV STIS spectra presented by Sonneborn et al. (1998) and Michael et al. (1998). In recent (January 2010) STIS observations, we tentatively

detected broad, redshifted N V emission (Figure 3 of France et al. 2010), but low S/N precluded a detailed analysis. Now, in our COS observations, we unambiguously detect this emission. While the blue wing of the line is lost under the bright Ly α emission, we observe the complete red wing of the N V λ 1240 velocity profile. In Figure 7, we compare the line profile of N V with that of He II, observing that the red wings of the two profiles are qualitatively similar. Both ions present a boxy line profile, with a flat top and fall off between +9000 – +10000 km s⁻¹. The N V profile extends to at least +14000 km s⁻¹, where the data are contaminated by the geocoronal O I triplet.

At the relevant energy scales of nonradiative supernova shocks, $v_s \gtrsim 10^3$ km s⁻¹, the excitation cross sections for Li-like species (e.g., C³⁺, N⁴⁺, O⁵⁺) by collisions with protons and heavier ions are considerably larger than the cross-section for ionization to their He-like stage (Laming et al. 1996; Borkowski et al. 1997). Therefore, these species may emit many (several hundred) line photons (predominantly C IV λ 1550, N V λ 1240, O VI λ 1032 for the ions listed above) for every atom that crosses the reverse shock front before becoming ionized. The total ratio of N V λ 1240 photons emitted to H α is given by

$$\frac{F(NV)}{F(H\alpha)} = \frac{1}{0.2} \frac{x_N}{x_H} \frac{R(1240)}{R(N^{4+} \rightarrow N^{5+})} \quad (2)$$

where 0.2 is the number of H α photons emitted per neutral hydrogen atom crossing the shock (Michael et al. 2003), x_N/x_H is the nitrogen abundance ratio, $R(1240)$ is the rate of electron and ion impacts that result in N V emission, and $R(N^{4+} \rightarrow N^{5+})$ is the rate of ionization to N⁵⁺. The relative nitrogen abundance in the circumstellar ring is $2.8 (\pm 1.1) \times 10^{-4}$ (Mattila et al. 2010). Given the similarity of the red side of the He II and N V profiles, we compute the total N V reverse shock photon flux by assuming that the relative red/blue contribution is the same as He II (Table 2) and interpolating through the narrow emission lines, as described above. Under these assumptions, the observed $F(NV)/F(H\alpha)$ ratio is ≈ 1.15 . Folding in the nitrogen abundance, we can therefore make a rough estimate of the ratio $R(1240)/R(N^{4+} \rightarrow N^{5+}) \sim 850_{-250}^{+550}$ required to explain the observed line ratio.

$R(1240)$ can be evaluated analytically,

$$R(1240) = n_e \langle \sigma_e v_s \rangle + n_H \langle \sigma_p v_s \rangle + n_{He} \langle \sigma_\alpha v_s \rangle \quad (3)$$

where the cross-sections, $\sigma_{e,p,\alpha}$, are for excitation of N V λ 1240 by electrons, protons, and alpha particles, respectively. We simplify the averages over velocity distribution by evaluating the expression at a single shock velocity, $v_s = 10^4$ km s⁻¹. Cross-sections for the proton and alpha particle collisions are taken from Table 2 of Laming et al. (1996) for proton energies ($E_p = 1/2 m_p v_s^2$) of 522 keV (we assume the 544 keV values) and alpha particle energies ($E_\alpha = 1/2 (4m_p) v_s^2$) of 2088 keV (we assume 1.2 times the 2720 keV values). The electron impact excitation cross-sections are not as well determined, but we expect them to be similar to those of the protons at high shock velocity. Therefore, we estimate $\sigma_{e,p,\alpha} = [4.31, 4.31, 14.4] \times 10^{-17}$ cm², respectively. Taking $n_e = 450$ cm⁻³ (Borkowski et al. 1997), $n_H \approx 100$

cm⁻³ (Smith et al. 2005; Heng et al. 2006), and $n_{He} = 0.17 n_H$ (Mattila et al. 2010), we compute a total N V excitation rate for $v_s = 10^4$ km s⁻¹; $R(1240) = 2.6 \times 10^{-5}$ s⁻¹.

Ionization rates by baryons are expected to be small compared to those by electrons due to the high energy of post-shock protons and heavier species (Laming et al. 1996), therefore we neglect their contribution to the N⁴⁺ ionization rates. Cross-sections for the direct ionization to N⁵⁺ by thermal electrons can be calculated as a function of electron energy from Equation 1 of Arnaud & Rothenflug (1985). Using coefficients for the Li-sequence (their Table 1), we calculate an electron ionization cross-section of $\sigma_{ion}^{NV} = 1.5 \times 10^{-18}$ cm² at $E_e = 107$ eV⁴. The direct ionization rate, $R(N^{4+} \rightarrow N^{5+}) = n_e \langle \sigma_{ion}^{NV} v_s \rangle = 1.5 \times 10^{-7}$ s⁻¹. This gives $R(1240)/R(N^{4+} \rightarrow N^{5+}) \approx 170$, which is outside the range allowed by the observations.

If, on the other hand, there is partial equilibration between the electron and proton distributions (Cargill & Papadopoulos 1988; Laming et al. 1996; Heng 2010), then $T_e' = \beta_{eq} T_p$, where

$$T_p = \frac{3}{16} \frac{\mu m_p v_s^2}{k_B} \quad (4)$$

and β_{eq} is the equilibration factor, which has a maximum value of 1. μ is the mean particle weight, $\mu \approx 0.55$ (for $x_H = 0.85$, $x_{He} = 0.14$, and $x_Z = 0.01$). In this case $E_e' = 1/2 m_e (2k_B T_e' / m_e) = k_B T_e'$. Using this energy to calculate the ionization cross-sections, we find that for $\beta_{eq} = 0.14 - 0.35$ ($T_e' = 1.7 - 4.4 \times 10^8$ K), the electron ionization cross-sections $\sigma_{ion}^{NV} = 9.7 - 4.2 \times 10^{-20}$ cm², and the ionization rates of $R(N^{4+} \rightarrow N^{5+}) = 4.3 - 1.9 \times 10^{-8}$ s⁻¹. This gives $R(1240)/R(N^{4+} \rightarrow N^{5+}) \sim 610 - 1380$, approximately the range required to account for the observed $F(NV)/F(H\alpha)$ ratio⁵. $v_s = 10^4$ km s⁻¹ is used as the fiducial velocity in our calculations, but the general conclusions hold for a range of possible shock velocities. Computing $R(1240)/R(N^{4+} \rightarrow N^{5+})$ for velocities $(5 - 12) \times 10^3$ km s⁻¹, we find that β_{eq} decreases by $\sim 50\%$ across this range, but $\beta_{eq} > 0.1$ is required to explain the data for all velocities considered.

It seems clear that partial ion-electron equilibration is required to explain the observation of strong reverse shock N V emission from SN 1987A. However, values of $\beta_{eq} = 0.14 - 0.35$ are greater than those favored by Laming et al. (1996) from their non-radiative shock modeling of SN1006. Moreover, equilibration values of $\beta_{eq} \geq 0.1$ for the high shock velocity ($v_s \sim 10^4$ km s⁻¹) in SN 1987A are inconsistent with some recent results on electron-ion equilibration in collisionless shocks. Ghavamian et al. (2007) discuss a sample of older ($\sim 10^3$ yr) supernova remnants whose electron-to-proton temperature ratio is inversely proportional to the square of the shock velocity, $\beta_{eq} \propto v_s^{-2}$; however, this relation

⁴ $E_e = \frac{3}{16} m_e v_s^2$

⁵ We note that if we assume that the ionization cross-sections scale as $\sigma_{ion}^{NV} \propto T_e'^{-0.5}$ in the high- T_e limit, then the ion-electron equilibration factor could be a factor of 2 – 3 higher, approaching the limit of complete equilibration.

has not been confirmed in subsequent work (Helder et al. 2011). Additionally, van Adelsberg et al. (2008) present evidence suggesting an increasing β_{eq} for $v_s > 2000$ km s⁻¹. Finally, we note that very high electron temperatures ($T_e \sim 10^9$ K) were inferred for the high velocity ejecta in SN 1993J (Fransson et al. 1996), suggesting that electron heating may be efficient in high velocity environments (see §5.3).

We also compare the observed N V flux with that predicted by Borkowski et al. (1997). Extrapolating their prediction to the time of our observations, we would have expected $(1.5 - 2.0) \times 10^{-13}$ erg cm⁻² s⁻¹ in 2011, or $0.06 - 0.08$ photons cm⁻² s⁻¹ after including a factor of 6.3 to account for interstellar reddening. We see that, similar to the prediction for He II, these values are a factor of ~ 2 below their observed 2011 values ($F(\text{N V}) = 0.16$ photons cm⁻² s⁻¹). Overall, we consider the prediction made 14 years prior to the observation to be remarkably good. The slight differences between the predicted and observed N V flux can most likely be attributed to their choice of a value for β_{eq} that is lower than we determine from the data. Lower electron energies increase the ionization cross-section and hence $R(N^{4+} \rightarrow N^{5+})$, thereby reducing the total number of N V photons emitted.

4.3.2. N IV] $\lambda 1486$

Figure 6 shows that there is considerable substructure on the continuum in addition to the strong reverse shock emission from H I, He II, and N V. The weakest of these features spans $\sim 1430 - 1500$ Å. We attribute this emission to N IV] $\lambda 1486$ from the reverse shock. The N IV] emission is highly asymmetric about the rest velocity. The blue-shifted component extends to roughly -8000 km s⁻¹, while the red side of the profile is at the noise level by $+3000$ km s⁻¹. Additionally, the blue wing of the reverse shock C IV profile, discussed in the next section, overwhelms any additional red flux. Again making the assumption that the red/blue N V ratio is the same as that for He II, and assuming that the red side of the N IV] profile contributes an additional 20% to the total flux, we find that the total $F(\text{N V})/F(\text{N IV])}$ ratio is $\sim 72 \pm 17$.

4.3.3. C IV $\lambda 1550$

Figure 6 also shows C IV $\lambda 1550$ from the reverse shock covering $\sim 1500 - 1580$ Å, with emission from $-12000 - -1500$ km s⁻¹ having 2.2 times the total integrated photon flux as the $+1500 - +8000$ km s⁻¹ interval (Table 2). The red side of the velocity profile is only marginally above the noise level at $v_{CIV} > +3000$ km s⁻¹. We consider the relative fluxes of C IV and N V here, and in the next section we will discuss the relative velocity distributions of the reverse shock species.

Interpolating the spectrum through the narrow emission lines, we measure an integrated broad line photon flux ($-12000 - +8000$ km s⁻¹) of $F(\text{C IV}) = 7.7 \times 10^{-3}$ photons cm⁻² s⁻¹. With the previously-noted assumptions about the N V flux in mind, we find a flux ratio of $F(\text{N V})/F(\text{C IV}) \approx 20 \pm 3$. The ratio of N V $\lambda 1240/\text{C IV}$

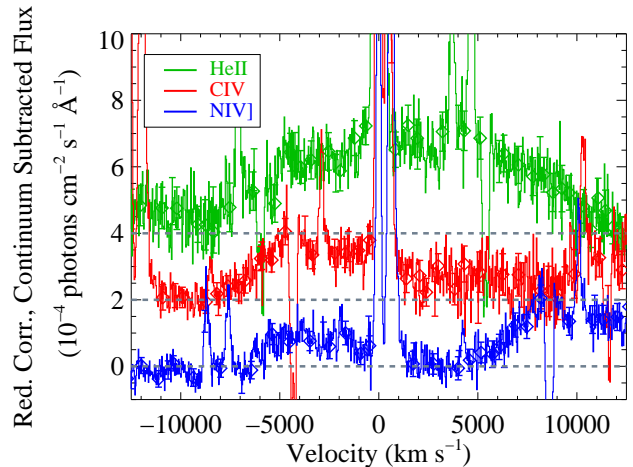


FIG. 8.— 2-photon continuum subtracted N IV], C IV, and He II in the combined WNW + ESE spectra. He II and C IV have been offset by 4.0 and 2.0 photons cm⁻² s⁻¹ Å⁻¹, respectively, for display purposes. The zero-flux levels for all three ions are shown as the dashed gray line.

$\lambda 1550$ is given by an expression similar to equation (2):

$$\frac{F(\text{NV})}{F(\text{CIV})} = \frac{x_N}{x_C} \left[\frac{R(1240)}{R(N^{4+} \rightarrow N^{5+})} \right] \left[\frac{R(1550)}{R(C^{3+} \rightarrow C^{4+})} \right]^{-1} \quad (5)$$

The ratio of the quantities in square brackets is very close to 0.9 for $v_s = (5 - 12) \times 10^3$ km s⁻¹ (the ratio is near unity because both are Li-like ions). That implies that the abundance ratio of nitrogen/carbon atoms crossing the reverse shock is $x_N/x_C \approx 22 \pm 3$, greater than the ratio $x_N/x_C \approx 8.5 \pm 3.5$ inferred from spectroscopic observations of the equatorial ring (Lundqvist & Fransson 1996; Mattila et al. 2010).

We note that while our estimate of the flux ratio N V $\lambda 1240/\text{C IV } \lambda 1550$ is somewhat uncertain because our choice of the blue side of the N V profile is speculative, the integrated flux in the red side of the N V profile alone is more than 10 times greater than the *total* C IV flux. Therefore, we are confident that the N V/C IV ratio is enhanced well beyond what can be attributed to the circumstellar ring abundances of the two species. We return to this point in Section 5.4.

5. DISCUSSION

5.1. Relative Velocity Distributions

In Sections 4.2 and 4.3, we describe the velocity distributions of He II, N V, N IV], and C IV. The red side of the N V velocity profile is compared with He II in Figure 7 and a comparison of the He II, N IV], and C IV velocity profiles is shown in Figure 8. We observe qualitatively different velocity profiles from He II, N IV], and C IV. He II displays a mildly asymmetric profile spanning $\sim -9000 - +11000$ km s⁻¹, while N IV] and C IV are strongly weighted towards the blue, spanning $\sim -(9000 - 8000) - +3000$ km s⁻¹. Peculiar geometric projection of the velocity can be ruled out because all three ions were observed with the G160M instrument mode, with a constant spacecraft roll orientation.

In nonradiative shocks, He II is thought to be excited

by collisions with electrons, while the Li-like species (N V and C IV) are excited by ion-collisions (Laming et al. 1996). As we have shown above, the electron and ion distributions have likely not equilibrated in the SN 1987A debris, therefore it may be that the relative distributions of the impacting species is reflected in the velocity profiles of the excited ions. One might also speculate that if ion collisions dominate the emission from metals, the preferentially blue N IV] and C IV profiles would reflect an asymmetric spatial distribution of the protons and ions. The N IV] and C IV profiles also show spatial variations between the WNW and ENE spectra. In contrast, the He II profile, which is primarily excited by electrons, is more spatially and spectrally symmetric. This may indicate a more uniform distribution of electrons. With these possibilities in mind, one is tempted to separate the metals from the helium, however we observe N V extending to red velocities as high as $+10000 \text{ km s}^{-1}$ before declining. Modeling of both the fluxes and velocity distributions of ions in the SN 1987A reverse shock would be very useful for our understanding of the observed velocity profiles.

5.2. Excess Emission at $\lambda < 1350 \text{ \AA}$

The high throughput and low instrumental background of COS have allowed us to study the ultraviolet continuum of SN 1987A for the first time. Similarly, we have presented the first conclusive observation of elements heavier than hydrogen in the reverse shock. Figures 3 and 6 show that there is additional ultraviolet continuum at $\lambda < 1350 \text{ \AA}$, above which we propose that hydrogen 2-photon emission can adequately explain the data. As we discuss in §4.1.2, the spectral contamination from Star 3 is inversely proportional to wavelength, as the cross-dispersion astigmatism height is largest at the shortest wavelengths. This could account for some, but probably not all, of the bluest continuum emission. Instead, we propose four speculative possibilities for this emission. He^+ has a $2p \rightarrow 1s \text{ Ly}\alpha$ line at 304 \AA , and will emit a $2s \rightarrow 1s$ 2-photon continuum which peaks at $\lambda \approx 355 \text{ \AA}$, and declines to the red. We estimate that this emission contributes $\approx 10\%$ of the 1150 \AA continuum.

There may be an additional $\text{Ly}\alpha$ component at very high velocities, although this would require $v_{\text{Ly}\alpha} \gtrsim 30 \times 10^3 \text{ km s}^{-1}$ ($\gtrsim 0.1 c$) to account for the reddest observed emission. Another possibility is that low-ionization line emission from the core, as opposed to higher-ionization emission from atoms crossing the reverse shock, contributes to the $\lambda \leq 1350 \text{ \AA}$ emission. Jerkstrand et al. (2011) describe 2-photon emission, scattered by low-ionization, high-opacity metal lines in the core, contributing to the observed spectrum at ~ 8 years following the explosion. The observed excess begins at roughly the wavelength of the O I] $\lambda 1356 \text{ \AA}$ emission line. If an appreciable amount of low-ionization metal emission escapes from the near side of the core, it may be responsible for a portion of the observed flux.

Perhaps the most likely scenario is one in which several minor reverse shock species contribute to the observed continuum between $1140 - 1350 \text{ \AA}$. We have seen that reverse shock emission from Be-like ions (e.g. NIV] $\lambda 1486$) are present. Therefore, emission from lines of

C III $\lambda 1175$ and O V $\lambda 1371$ may also contribute⁶.

5.3. Partial Ion-Electron Equilibration in SN 1987A: Evidence for the Cross-Shock Potential?

In Section 4.3.1, we found that large ion-electron temperature equilibration ratios ($\beta_{eq} \approx 0.14 - 0.35$) are required to explain the high N V $\lambda 1240$ flux observed in the spectrum of SN 1987A. Ghavamian et al. (2007) predict equilibration ratios several times less than this for shock velocities observed in SN 1987A ($v_s \sim 10^4 \text{ km s}^{-1}$). Therefore, we consider additional electron heating mechanisms that could explain the observed N V emission. One mechanism that has been proposed for collisionless shocks (usually assumed to be quasi-perpendicular) at low Alfvén mach numbers (M_A) is the cross-shock potential (e.g Balikhin et al. 1993). But the cross-shock potential is expected to become less important as M_A increases because the plasma resistivity causes insufficient dissipation and the shock becomes unsteady (Edmiston & Kennel 1984). Electrons undergoing $\vec{E} \times \vec{B}$ drift across the shock front are not guaranteed to see the same cross-shock potential. They may lose energy as well as gain it, and this limits the electron heating.

However, it has also been suggested that at Alfvén Mach numbers appropriate to shocks in supernova remnants, the shock transition becomes thin. The shock transition can develop a length scale less than the electron convective gyroradius, thus eliminating any $\vec{E} \times \vec{B}$ drift within the shock. This onsets at $M_A = \sqrt{m_i/m_e}$, and electron heating by the cross shock potential again becomes viable (Balikhin et al. 1993; Gedalin et al. 2008). Typical SNR shocks have M_A much higher than this, neglecting any modification of the preshock medium. Yet the predicted heating is not observed. Most likely, M_A at these shocks is much lower than expected due to amplification of the magnetic field by cosmic rays streaming ahead of the shock (Bell 2004, 2005). Electron heating in these systems occurs via plasma waves excited upstream by shock reflected ions or cosmic rays. At higher velocity shocks, or for those likely to have weak upstream magnetic field due to their environment and/or insignificant cosmic ray populations, $M_A > \sqrt{m_i/m_e}$ and the cross-shock potential may be capable of heating electrons to the required energy. Therefore the energetic environment of SN 1987A, especially at its reverse shock, may enable additional electron heating mechanisms that do not contribute in lower M_A systems.

5.4. The N V/C IV Ratio

The final unresolved issue is the large N V/C IV ratio in the reverse shock emission. Carbon and nitrogen abundance ratios in the circumstellar ring suggest this number should be ≈ 8 (Lundqvist & Fransson 1996; Mattila et al. 2010), however the observed flux ratios indicate $\text{N/C} \approx 22$. It may be that the N/C abundance ratio in the outer envelope of the progenitor was stratified prior to the ejection of the circumstellar rings, and that we are now seeing the first observational evidence of that stratification. A second possibility to

⁶ O V is the next ionization stage down from the Li-like O VI $\lambda 1032$ line which we do not detect, but is presumably in the reverse shock spectrum at shorter wavelengths.

account for this discrepancy is that ongoing thermonuclear processing continued to convert C to N in the supernova progenitor following the circumstellar ring ejection. The CNO bi-cycle will enrich the ^{14}N abundance at the expense of the abundances of ^{12}C and ^{16}O and, in equilibrium, will convert almost all of the primordial C and O into N (Caughlan & Fowler 1962). CNO processing has been invoked to explain the fact that the observed nitrogen abundance in the circumstellar ring is elevated by a factor ~ 10 over its value in the LMC (Fransson et al. 1989). Our observation that the He abundance does not change between the ring and reverse shock is qualitatively consistent with the 14E1 model presented by Shigeyama & Nomoto (1990). They show that the high-velocity material we observe crossing the reverse shock front is only a small fraction of the total ejected mass, and does not probe deep enough into the interior of the ejecta to observe significantly elevated He abundances.

If CNO processing continued near the stellar surface following the ejection of the circumstellar rings, it could have in principle converted most of the remaining C and O abundances seen in the equatorial ring into N. Heng et al. (2008) note reduced oxygen abundances, possibly related to the high N V/C IV ratio observed in our observations. For this explanation to be viable, the timescale to reach equilibrium in the CNO cycle must be $\lesssim 20,000$ years, the interval since the ejection of the equatorial ring. This condition will be met if the temperature of the shell where CNO burning takes place is $T \gtrsim 3.5 \times 10^7$ K (Caughlan and Fowler 1962, Table 4). This effect could not only account for the high N/C ratio in the gas crossing the reverse shock, it could also increase the N/H ratio by an additional factor of ~ 2 . This would lower the required β_{eq} derived in §4.3.1.

6. CONCLUSIONS

We have presented an analysis of deep *HST*-COS spectroscopy of SN 1987A. Below, we summarize our primary results.

1. We observe narrow lines from shocked gas in the circumstellar ring, broad emission lines from the reverse shock, and a strong detection of ultraviolet continuum emission. Several of the emission line species and the far-UV continuum are conclusively detected for the first time.
2. The asymmetry in $\text{Ly}\alpha$ profile and its enhancement relative to $\text{H}\alpha$ suggests that most of the $\text{Ly}\alpha$ emission is a result of the illumination of the outer supernova debris by soft ($E \leq 0.5$ keV) X-rays emitted by the shocked circumstellar ring.
3. The $\lambda \gtrsim 1350$ Å continuum may be described by hydrogen 2-photon emission originating in the outer ejecta, near the reverse shock front.
4. We present resolved velocity profiles of He II, C IV, N IV], and N V from the reverse shock. A comparison of the velocity integrated He II and $\text{H}\alpha$ velocity profile indicates a He/H abundance ratio of 0.14 ± 0.06 . In order to reproduce the observed $F(\text{N V})/F(\text{H}\alpha)$ line ratio, partial ion-electron equilibration is required, $T'_e = \beta_{eq}T_p$, with $\beta_{eq} \approx 0.14 - 0.35$. Large values of β_{eq} may be explained by additional electron heating by the cross-shock potential.
5. The velocity profiles of C IV and N IV] are significantly different from that of He II, which may be related to the different excitation processes (ion vs. electron collisions) for the different species.
6. We observe additional continuum emission at $\lambda < 1350$ Å that is not readily explained. We favor a scenario where several weaker emission lines contribute, but spectral overlap with the much stronger $\text{Ly}\alpha$ and N V emission profiles prevents an unambiguous interpretation.
7. Finally, we observe that the C IV reverse shock emission is weaker than would be expected assuming circumstellar ring abundances. This may be explained by chemical stratification in the outer envelope of the progenitor, and may indicate additional CNO processing between the period of circumstellar ring ejection and the supernova explosion.

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