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Core Evolution Driven by Mantle Global Circulation

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Core evolution driven by mantle global circulation

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Abstract

Reconstructions of the Phanerozoic history of mantle global circulation that include past plate motions are used to constrain the thermochemical evolution of the core. According to our mantle global circulation models, the present-day global average heat flux at the core-mantle boundary lies in the range 80-90 $\mathrm{mW.m^{-2}}$, with peak-to-peak, long wavelength lateral variations up to 100 $\mathrm{mW.m^{-2}}$ associated with compositional and thermal heterogeneity in the D"-layer. For core thermal conductivity in the range k=100-130 $\mathrm{W.m^{-1}.K^{-1}}$ we infer that the present-day outer core is thermally unstable beneath the high seismic velocity regions in the lower mantle but thermally stable beneath the large low seismic velocity provinces. A numerical dynamo shows how this boundary heat flux heterogeneity generates departures from axial symmetry in the time average geomagnetic field and the pattern of flow in the outer core. Standard thermochemical evolution models of the core driven by mantle global circulation heat flow predict inner core nucleation between 400 and 1100 Ma. With thermal conductivity $k \simeq 100$ $\mathrm{W.m^{-1}.K^{-1}}$ the core heat flow derived from our mantle global circulation models is adequate for maintaining the geodynamo since inner core nucleation, supercritical for dynamo action by thermal convection just prior to inner core nucleation, and marginal for inner core convection.

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1 Introduction

The geodynamo owes its existence to convection in the mantle. The rate of energy release required to maintain the geodynamo at its present-day intensity over geologic time is so large – on the order of 10-16 TW (terawatts) – that it would likely have ceased to operate long ago were it not for the heat extracted from the core by the circulation of the mantle. Estimates of the energy required by the geodynamo as well as estimates of the actual heat loss from the core have recently been revised upward, partly in response to recent studies indicating the thermal conductivity of core alloys is higher than previously assumed (de Koker et al., 2012; Pozzo et al., 2012; Gomi et al., 2013; Zhang et al., 2015), and partly because the radial structure and the amount of lateral heterogeneity in the D" region near the base of the mantle imply that the heat flow from the core is large (Buffett, 2007; Hernlund, 2010; Zhang and Zhong, 2011; Wu et al., 2011).

The combination of higher thermal conductivity and high core heat flow implies that the rate at which the core evolves is also fast in comparison with what would be the case were these properties smaller. An often-used metric for core evolution is the rate of growth of the solid inner core. Assuming the inner core boundary is at the melting point and the outer core is well mixed, growth of the inner core by solidification must track the cooling of the core as a whole (Labrosse, 2003; Buffett, 2003). In addition, the inner core growth contributes directly to maintaining the geodynamo through release of buoyant lighter elements, driving thermochemical convection in the liquid outer core (Jones, 2007).

Major problems for quantifying the energy budget of the core and its rate of evolution stem from uncertainties in the core-mantle boundary (CMB) heat flow, the melting curve in the core (Andrault et al., 2011; Anzellini et al., 2013), the partitioning of light elements at the inner core boundary (Gubbins et al., 2004; Nimmo and Alfe, 2006), and the amount of radioactive heat production in the core (Gessmann and Wood, 2002; Murthy et al., 2003; Bouhifid et al., 2007; Hirose et al., 2013). Among these parameters, the CMB heat flow is probably the most important and is certainly the most complex, because the local heat flux is inhomogeneous on the CMB and the total heat flow from the core varies with time.

All estimates of the present-day core heat flow are all based on indirect methods; these include calculation of mantle plume fluxes, consideration of dynamo thermodynamics, interpretations of lower mantle seismic structure, and output from mantle global circulation models (hereafter referred to as mantle GCMs). Mantle plume flux calculations based on hotspot activity initially yielded small values, in the range of Q_{cmb}= 2-5 TW (Loper, 1978; Davies, 1988; Sleep, 1990; Stacey and Loper, 2007) for the total core-mantle boundary heat flow, although later improvements to these estimates (Labrosse, 2002) yielded $Q_{cmb} \simeq 13$ TW (Leng and Zhong, 2008). Estimates derived from the thermodynamics of the geodynamo yield somewhat higher values, generally in the range Q_{cmb}= 4-10 TW (Buffett et al., 1996; Buffett, 2002; Labrosse, 2003; Gubbins et al., 2004). Interpretations of the seismic structure in the D" region at the base of the mantle in terms of post-perovskite phase changes yield significantly higher values, with average heat flux in the range $\bar{q}_{cmb}$=65-100 mW.m^{-2} (Lay et al., 2006; van der Hilst et al., 2007; Monnereau and Yuen, 2010; Wu et al., 2011) equivalent to a total core heat flow of Q_{cmb} =10-16 TW, although Tateno et al. (2009) obtained Q_{cmb}=6 TW with this approach. Interpretations of the lateral heterogeneity in the seismic structure also provide estimates of the lateral heterogeneity in CMB heat flux in the range of q'_{cmb}= 20-50 mW.m^{-2} (van der Hilst et al., 2007; Lay et al., 2008). Not surprisingly, such a wide range of the core heat flows yield a comparably wide range for the age of inner core nucleation, hereafter abbreviated ICN. The lower core heat flow estimates predict ICN ages in excess of 2.5 Ga, whereas the higher estimates predict ICN ages around 0.5 Ga (Labrosse et al., 2001; Roberts et al., 2003; Nimmo, 2007). Adding to this uncertainty, the CMB heat flow is time dependent, yet there is little by way of direct observational constraints on how much it has varied since the ICN.

Dynamically based predictions for the time variation of the average core heat flow and its lateral heterogeneity can be extracted from mantle GCMs. The CMB heat flow in these models depends on many parameters, including the lower mantle viscosity, thermal conductivity, and the thermal gradient in the D" region, the latter depending on the strength of the circulation in the lower mantle, the compositional stratification, phase changes in D" such as post-perovskite (Nakagawa and Tackley, 2011), and the presence or absence of smaller scale instabilities in that region (Nakagawa and Tackley, 2010; Zhang et al., 2010; Zhang and Zhong, 2011). Uncertainties in these mantle properties, as well as the non-uniqueness in the surface plate reconstructions that are often used as upper boundary conditions lead to substantial uncertainty in mantle GCM predictions.

However, mantle GCMs can be tuned to match the present-day surface heat flow and can also be tuned to match the present-day internal structure of the mantle, reducing their uncertainty somewhat. In this connection, the structure of dense chemical piles in the lower mantle offers an important geodynamical constraint on core heat loss. It is found that very high CMB heat flow is required to maintain compositionally dense piles the size of the two large low shear velocity provinces (LLSVPs) seen in the present-day lower mantle seismic structure (McNamara and Zhong, 2004). Depending on the values of other mantle parameters, maintaining two dense piles comparable in size to the LLSVPs requires a mean CMB heat flux of $\bar{q}_{cmb}$=75-100 mW.m^{-2} and peak-to-peak, long wavelength lateral variations up to 100 mW.m^{-2} (Nakagawa and Tackley, 2008; Zhang and Zhong, 2011; Olson et al., 2013).

In this paper we use statistics of the global mean CMB heat flow and lateral variations of CMB heat flux obtained from plate-driven mantle GCMs that generate lower mantle chemical piles similar to those observed in the lower mantle to calculate the thermal evolution of the core backward in time, starting from the present-day and continuing to the time of ICN. We also use the present-day pattern and magnitude of CMB heat flux from one of these mantle GCMs to drive a numerical dynamo model, linking the structure of the dynamo-produced magnetic field and lateral heterogeneity within the outer core to the global mantle circulation.

2 Mantle global circulation and core heat flux

Mantle global circulation models provide self-consistent relationships between dynamical properties of the mantle such as plate spreading rates, viscosity, and radioactive heat production and core heat flux, and observables such as mantle heterogeneity and heat flux at the surface (McNamara and Zhong, 2005). In some mantle GCMs the circulation is entirely free convection driven by thermal and compositional buoyancy (Nakagawa and Tackley, 2013, 2014). In others, the circulation is a combination of forced convection driven by prescribed surface plate motions plus free convection (McNamara and Zhong 2005; Zhang et al., 2010; Zhang and Zhong, 2011; Bower et al., 2013; Bull et al., 2014; Rudolph and Zhong, 2014). A commonly-used procedure in these models is to adjust the Rayleigh number governing the free convection part of the circulation to match some global constraint, such as zero net torque on the surface plates or equal r.m.s. velocity of the free and forced components of the flow.

Table 1 gives the input parameters of the mantle GCM used in this study. In addition to transport and thermodynamic parameters, the mantle GCM depends on the prescribed surface plate motions. Here we have used four paleoplate reconstructions. Case 1 uses the reconstruction by Muller et al. (2008) covering the period 0-140 Ma; Case 2 uses the reconstruction by Lithgow-Bertelloni and Richards (1998) covering the period 0-119 Ma; Case 3 uses the reconstruction by Seton et al. (2012) covering 0-200 Ma. Each case has identical initial conditions, including an initially 250 km thick dense layer at the base of the mantle, with properties listed in Table 1. Each case started at 608 Ma, with the first 150 Myr as a spin-up phase. The spin-up phase was initiated using a horizontally uniform temperature field taken from a pre-calculation run to statistically steady state with rms surface velocity chosen to match the rms velocity of the first (450 Ma) stage of the Zhang et al. (2010) 450-119 Ma proxy plate reconstruction. Our Case 2 is identical to the reference case FS1 in Zhang et al. (2010) and to Case HF1 from Zhang and Zhong (2011). It is also the same as Case 2 in Rudolph and Zhong (2014). Our Cases 1 and 3 are identical to our Case 2 except for the plate motions over the last 200 Ma in our Case 3 and over the last 140 Ma in our Case 1, for which Seton et al. (2012) and Muller et al. (2008) are used, respectively.

We use temperature-dependent viscosity η with a depth-dependent viscosity prefactor of the form

$$\eta = \eta_0 \exp\left(E^*(0.5 - T^*)\right) \tag{1}$$

where η_0 is a depth-dependent viscosity prefactor, E^* controls temperature-dependence and T^* is non-dimensional temperature, which varies from 0 at the surface to 1 at the CMB. We use E^* =9.21, leading to variations in viscosity of four orders of magnitude from temperature variations. We include a 30-fold decrease in viscosity prefactor at 150 km depth, a uniform viscosity prefactor in the upper mantle and transition zone, a factor of 60 increase in viscosity prefactor at 670 km depth, and a linear increase in viscosity prefactor across the lower mantle leading to an overall factor of 3.4 increase. This viscosity structure is identical to that used in Rudolph and Zhong (2014) Case 2, Zhang et al. (2010) Case FS1, and Zhang and Zhong (2011) Case HF1. We use a numerical resolution of 64^3 elements on each of the 12 caps of the CitcomS mesh with refinement in the radial direction in boundary layers.

Figure 1 shows the variation in the global average CMB heat flux $\bar{q}_{cmb}$ versus age from three mantle GCMs calculated using three plate tectonic reconstructions as surface boundary conditions. Figure 1 also shows heat flux patterns on the CMB at four distinct times in the Phanerozoic from mantle GCM Case 2. The continent locations are shown in shadow, and convergent and divergent plate boundaries are shown by solid and dashed lines, respectively (Zhang et al., 2010). These images represent the longest-wavelength components of the CMB heat flux heterogeneity, represented by spherical harmonic degrees 1-4.

Several points are worth noting here. First, the present-day CMB heat flux pattern in Figure 1a is dominated by the spherical harmonic degree 2 structure that is prominent in lower mantle seismic tomography (Romanowicz and Gung, 2002; Dziewonski et al., 2010; Lekic et al. 2012). High heat flux is distributed along an approximately great circle band passing beneath the eastern parts of the Americas and Asia. Low heat flux occurs in two regions, one beneath Africa the other beneath the central Pacific, closely coincident with the seismically observed LLSVPs. In terms of the dynamics of the lower mantle, the high CMB heat flux belt corresponds to lower mantle downwellings where lithospheric slabs descend toward the CMB; the low CMB heat flux regions correspond to lower mantle upwellings above the dense chemical piles, which have been implicated as sites of deep mantle plume formation (Burke and Torsvik, 2004; Burke et al., 2008; Torsvik et al., 2006). In contrast, at 275 Ma in Figure 1 the CMB heat flux is dominated by a spherical harmonic degree 1 pattern, with mostly high heat flux beneath the margins of supercontinent Pangaea produced by major downwellings originating at convergent plate margins arrayed around the edge of the supercontinent. This spherical harmonic degree 1 pattern is partially disrupted around 180 Ma by the breakup of Pangaea and is further disrupted by opening of the Atlantic, so that by 110 Ma the CMB heat flux pattern is dominated by a spherical harmonic degree 2 very similar to the present-day.

The present-day global mean CMB heat flux in Figure 1b is $\bar{q}_{cmb} = 86$ mW.m^{-2}, less than the $q_{ad} \simeq 100$ mW.m^{-2} conducted down the core adiabatic gradient if we assume a high value of $k = 130$ W.m^{-1}.K^{-1}for the thermal conductivity in the outer core below the CMB (corresponding to about 15 TW total core heat flow). The difference between the global mean CMB heat flux and adiabatic conduction suggests the presence of stable thermal stratification in the outer core beneath the CMB, with the possibility that thermal convection might be suppressed there. However, it is necessary to take into account the lateral heterogeneity in CMB heat flux produced by the lower mantle convection. The hatched contours in Figure 1 enclose regions where the local CMB heat flux q_{cmb} exceeds 100 mW.m^{-2}; these regions cover nearly 40% of the CMB at the present-day, nearly 45% at 110 Ma, and about 30% at 275 Ma, respectively. Within these regions the local CMB heat flux is expected to exceed the heat conducted down the outer core adiabat even if the thermal conductivity of the outer core is as high as 130 W.m^{-1}.K^{-1}. The reverse situation applies in regions outside the hatched contours; there we expect stable thermal stratification beneath the CMB if the thermal conductivity is high. Whether or not such a patchwork of superadiabatic and subadiabatic heat flux supports a global layer with stable stratification beneath the CMB remains an open question. Buffett (2014) has interpreted the geomagnetic secular variation in favor of global thermal stratification beneath the CMB, whereas Amit (2014) came to the opposite conclusion using the same data. Another possibility is compositional stratification due to light element gradients in this region (Helffrich and Kaneshima, 2010), which could be far more stabilizing than purely thermal stratification.

3 Heterogeneous core-mantle boundary heat flux and the present-day geodynamo

We model the influence of the general circulation of the mantle on the present-day state of the geodynamo by applying the CMB heat flux pattern shown in Figure 1a to a numerical dynamo driven by the coupled effects of CMB heat flux and chemical differentiation at the inner core boundary associated with inner core growth. The standard approach to modeling Boussinesq thermochemical convection in the outer core involves the co-density variable

$$C = \rho_{oc} \left(\alpha T + \beta \chi \right) \tag{2}$$

where ρ_{oc} is average outer core density, T is the outer core temperature relative to the adiabat, χ is the outer core light element concentration, and α and β are volumetric expansivities for T and χ, respectively. At the CMB we specify the heat flux as the sum of a global mean part $\bar{q}_{cmb}$ and a laterally varying part q'_{cmb}:

$$q_{cmb} = \bar{q}_{cmb} + q'_{cmb} \left(\phi, \theta \right) \tag{3}$$

where ϕ and θ are longitude and co-latitude, respectively. $\bar{q}_{cmb}$ is to be compared with the heat conducted down the core adiabat q_{ad}, such that $\bar{q}_{cmb} - q_{ad} > 0$ corresponds to superadiabatic heat flux in the Boussinesq approximation. The function q'_{cmb} in (3) specifies the amplitude and the planform of the CMB heat flux heterogeneity.

Writing the codensity as the sum of global mean and laterally varying parts $C = \bar{C} + C'$, we express the CMB heat flux (3) as

$$\left.\frac{\partial \bar{C}}{\partial r}\right|_{\text{cmb}} = -\frac{\rho_{oc}\alpha(\bar{q}_{cmb} - q_{ad})}{k}; \qquad \left.\frac{\partial C'}{\partial r}\right|_{\text{cmb}} = -\frac{\rho_{oc}\alpha q'_{cmb}}{k} \tag{4}$$

where k is the outer core thermal conductivity. At the inner core boundary (ICB) we assume constant codensity $C = C_{icb}$.

We take $\bar{q}_{cmb}$ and q'_{cmb} from Figure 1a and convert these to codensity boundary conditions using (4). We nondimensionalize these boundary conditions for input into the numerical dynamo using the difference between CMB and ICB radii $D = r_{cmb} - r_{icb}$ and D^2/ν to scale length and time, respectively, and $\rho_{oc}\beta D^2\dot{\chi}/\nu$ to scale co-density, where ν is outer core kinematic viscosity and $\dot{\chi}$ is the time rate of change of the light element concentration in the outer core due to inner core growth, which is the main source of buoyancy for outer core convection. This choice of scaling produces the following dynamo control parameters (Olson et al., 2013): the compositional Rayleigh number and Ekman number

$$Ra = \frac{\beta g D^5 \dot{\chi}}{\kappa \nu^2}; \qquad E = \frac{\nu}{\Omega D^2} \tag{5}$$

where g is gravity at the CMB and Ω is the angular velocity of rotation, plus the Prandtl and magnetic Prandtl numbers

$$Pr = \frac{\nu}{\kappa}; \qquad Pm = \frac{\nu}{\eta} \tag{6}$$

where κ is diffusivity for the codensity. The heat flux boundary conditions at the CMB (4) are given in terms of the dimensionless codensity (denoted with asterisks) as

$$\left.\frac{\partial \bar{C}*}{\partial r*}\right|_{\text{cmb}} = -\frac{Ra_q}{Ra}; \qquad \left.\frac{\partial C'*}{\partial r*}\right|_{\text{cmb}} = -\frac{Ra_{q'}}{Ra} f^* \tag{7}$$

where the Rayleigh numbers based on CMB heat flux are defined as

$$Ra_q = \frac{\alpha g D^4 (\bar{q}_{cmb} - q_{ad})}{\nu \kappa k}; \qquad Ra_{q'} = \frac{\alpha g D^4 \delta q_{cmb}}{\nu \kappa k} \tag{8}$$

with δq_{cmb} = max(q'_{cmb})- min(q'_{cmb}) and f^*=$q'_{cmb}/\delta q_{cmb}$.

Figures 2, 3, and 4 show snapshots and time averages of the structure of a thermochemical numerical dynamo defined according to (2-8) with Rayleigh number $Ra = 4 \times 10^6$, Ekman number $E = 10^{-4}$, Prandtl number Pr =1, magnetic Prandtl number Pm =6, ϵ = -1.47 for the codensity sink (see Supplementary Materials), plus the CMB heat flux from Figure 1a with Ra_q/Ra = - 0.08 and $Ra_{q'}/Ra = 0.1$, corresponding to an assumed $q_{ad} = 100$ mW.m^{-2} from k=130 W.m^{-1}.K^{-1}in the outer core. The numerical dynamo code (MagIC; Wicht, 2002) used 81, 128, and 256 outer core grid points in radius, latitude, and longitude, respectively, 9 radial points in the inner core, and spherical harmonic truncation at degree and order 85. Time averages were computed over 10 magnetic dipole diffusion times, corresponding to roughly 500 kyr in the core. No polarity reversals were recorded.

Figure 2 shows a snapshot of the radial component of the magnetic field on the CMB from the numerical dynamo, compared with the radial component of the modern geomagnetic field on the CMB from core field model POMME 2008 truncated at spherical harmonic degree and order 12. Contours of the geomagnetic field are in millitesla; contours of the dynamo field are in dimensionless Elsasser number units $\sigma B^2/\rho_{oc}\Omega$, where σ is the electrical conductivity of the core and B is the magnetic field intensity. Magnetic structures that are suggestive of the modern core field include the high intensity flux lobes under North America and Eurasia, the longitudinal strip of intense field beneath Australia, and subequatorial patches of reversed flux that drift westward, which in the dynamo are advected by east-to-west azimuthal flow. These magnetic structures, particularly the high latitude patches, represent the tops of quasi-columnar convective structures extending deep into the outer core that become amplified by downwelling flow as they pass beneath regions with high CMB heat flux.

The effects of the CMB heterogeneity can be seen in the deviations from axisymmetry in the time average CMB magnetic field shown in Figure 3a, including higher intensity field lobes in the northern hemisphere at the longitudes where the CMB heat flux is maximum. Reduced versions of these lobes are also evident in the southern hemisphere, but there the non-axisymmetric structure merges into a single high latitude lobe, as found previously in dynamos using tomographic CMB heat flux conditions (Olson and Christensen, 2002). The radial velocity pattern in Figure 3b shows departures from axial symmetry induced by the CMB heterogeneity, particularly beneath Asia, superimposed on the stronger downwelling induced by the inner core tangent cylinder.

CMB heat flux heterogeneity is felt all the way to the ICB. Figure 3c shows the time average of the codensity flux on the ICB, contoured such that red corresponds to the largest flux and blue to the smallest. According to the definition (2), lateral variations in ICB codensity flux in this dynamo can be considered as a proxy for the lateral variations in the rate of inner core solidification. The large zonal variation in Figure 3c, with high codensity flux at low latitudes and low codensity flux at high latitudes is characteristic of the heat and light element fluxes produced by the columnar structure of the convection, which advects the codensity more efficiently outside the inner core tangent cylinder. However, the nonzonal variations in Figure 3c are products of the CMB heterogeneity. In addition to a spherical harmonic degree 2 modulation there is also a spherical harmonic degree 1 component, marked by a low latitude concentration of codensity flux with its maximum located in the Eastern hemisphere. This transformation of dominantly spherical harmonic degree 2 CMB heterogeneity into spherical harmonic degree 1 ICB heterogeneity by the flow in the outer core has been found previously in numerical dynamos (Aubert et al., 2008) and has been suggested as a driver for the hemispherical differences observed in the seismic structure of the inner core.

Additional effects of the CMB heterogeneity are evident in the time average codensity structure shown in Figure 4. The deviations from azimuthal symmetry in Figure 4a, most evident in the region just below the CMB, are consequences of the lateral variations in CMB heat flux producing radial downflows while attenuating azimuthal motion at longitude bands where the CMB heat flux is highest, and producing radial upflows while enhancing azimuthal motion at longitude bands between these. The equatorial mean codensity profile in Figure 4b includes a thin layer just below the CMB in which the codensity gradient is slightly positive and therefore stable, a consequence of the equatorial mean CMB heat flux being subadiabatic. Although the stratification is locally stable, especially beneath the low CMB heat flux regions, the average stratification is practically neutral, as the global mean profile in Figure 4b demonstrates. We find that this type of patchwork stratification has little effect on the overall behavior of the dynamo. For example, Figure 3 shows that weak radial motions penetrate close to the CMB in many places in spite of the patchwork stratification. We note that these weak upwellings and downwellings are nevertheless strong enough to produce magnetic flux concentrations on the CMB that are morphologically similar to the flux concentrations in the present-day core field in Figure 2 and also appear in the time averaged core field (Johnson and Constable, 1995).

The structure of this dynamo would likely be different had we imposed stratification on the outer core, rather than allow stratification to develop from an initially adiabatic core as a consequence of the competition between positive and negative buoyancy fluxes originating at the ICB and CMB. Imposed stratification can be made arbitrarily strong, dividing the outer core convection into distinct layers for example (Nakagawa, 2011). With our method, stratification is dynamically limited by the magnitude of the stabilizing boundary flux, which in our case is relatively small.

4 Mantle-driven evolution of the core

The three mantle GCMs in Figure 1 show the same general trends in mean CMB heat flux with time. In each case the global mean CMB heat flux rises to $\bar{q}_{cmb} = 85$ mW.m^{-2} near 220 Ma, then peaks at 88-94 mW.m^{-2} around 70 Ma, before falling to 81-86 mW.m^{-2} at present. The minor differences in $\bar{q}_{cmb}$ prior to 220 Ma are numerical, attributable to differences in the precision of the tracer methods that are used to track the compositional heterogeneity in the three cases. Overall, the variation between the three cases is generally smaller than the peak-to-peak variation within a single case. For these three cases the mean and standard deviation of the 0-200 Ma total core heat flow correspond to $Q_{cmb} = 13.1 \pm 1.3$ TW. As discussed earlier, the CMB heat flow in mantle GCMs depends on the temperature on the CMB as well as transport properties in the mantle, particularly mantle viscosity and thermal conductivity. Other mantle GCMs by Zhang and Zhong (2011) examined the effects on CMB heat flow due to absence of the D" chemical layer, differences in mantle viscosity structure, changes in the Clapyeron slope of the transition zone phase transformations, as well as increase in the spreading rate of the Pacific oceanic plates. Varying these parameters yielded time average CMB heat fluxes generally higher than the preferred case, spanning the range 80-110 mW.m^{-2}, or approximately 12-17 TW. Similarly, Wu et al. (2011) obtained $Q_{cmb} = 13 \pm 3$ TW in their inversion of mantle lower mantle tomographic structure. Accordingly, for calculating the evolution of the core we focus on the range Q_{cmb}=12-14 TW as being representative of the past few hundred million years, but we consider cases in which Q_{cmb} deviates from this range by as much as ± 4 TW. This covers the spread of core heat flow produced by other mantle GCMs that support chemical piles in the D"-layer (Nakagawa and Tackley, 2005, 2013; Zhang and Zhong, 2011).

The dynamo model results in the previous section demonstrate that the CMB heat flow predicted by mantle GCMs, although comparable to or slightly less than adiabatic, can produce dynamo magnetic field structures similar to what is observed in the present-day core field, provided no strong compositional layering is present. Obvious follow-up questions are: what are the implications for this state of the core going backward into the deep past? For how long is this thermal regime viable in terms of its ability to maintain the dynamo, and similarly, what is the age of the inner core implied by this thermal regime?

Figure 5 shows how the evolution of the core is modeled since the time of ICN. The solid curves represent the present-day adiabatic temperature profile T_{ad} and light element concentration χ, and the dotted curves are the same at the time of ICN. The dashed red curve is the melting curve in the core denoted by T_{melt}, the total heat loss from the core to the mantle at the CMB is denoted by Q_{cmb}, and the total heat production within the core by radioactive decay is denoted by Q_{rad}. In calculating the evolution of the core it is usually assumed that the inner core boundary is a phase equilibrium boundary between the solid inner core and the liquid outer core so that $T_{icb} = T_{melt}$ at r_{icb}, the radius of the ICB. We also assume, consistent with the results of our numerical dynamo, that the outer core is well-mixed and therefore the geotherm closely follows an adiabatic temperature profile T_{ad}, the light element concentration in the outer core is uniform, and that the adjustment time of the dynamics in the core is small compared to the timescale for changes in the thermal structure of the lower mantle and core, so that the outer core remains in a state of statistical thermal and compositional equilibrium with respect to $Q_{cmb} - Q_{rad}$ (Buffett et al., 1996; Nimmo, 2007).

With these assumptions, the rate of inner core growth in response to the cooling of the core can be written (Labrosse, 2003)

$$\dot{r}_{icb} = \frac{(Q_{cmb} - Q_{rad})}{P} \qquad (9)$$

where the $P = P_l + P_g + P_s$ is the sum of individual contributions to the core energy balance from latent heat release at the ICB, gravitational energy release, and secular cooling of the core, respectively. Expressions for the individual contributions to P are given in the Supplementary Materials section in terms of core properties. Overall, P is most sensitive to the difference between the gradients of the core adiabat T_{ad} and the melting curve T_{melt} at the ICB, i.e., the parameter

$$\Theta = \left(\frac{dT_{ad}}{dr} - \frac{dT_{melt}}{dr}\right)\Big|_{\text{icb}} \qquad (10)$$

As shown in Figure 5, the combination of large $Q_{cmb} - Q_{rad}$ and small Θ implies relatively fast inner core growth, whereas the combination of small $Q_{cmb} - Q_{rad}$ and large Θ implies relatively slow inner core growth.

Our procedure for calculating the evolution of the core and the inner core age consists of the following steps: We first define a range of CMB heat flow based on the mantle GCMs described above. Next, we backward integrate (9) starting from the present-day, tracking the evolution of the core to determine the ICN age, examining the widest plausible ranges of Q_{cmb}, Q_{rad}, and Θ, the latter calculated by varying the assumed melting temperature at the ICB, $T_{melt}(r_{icb})$, away from its nominal value given in Table 2. Finally, we test the viability of the geodynamo across this parameter range by calculating from dynamo scaling laws the magnetic Reynolds number of outer core convection, to assess whether the core evolution model is consistent with maintaining the geodynamo both after and before ICN.

Implicit in the above procedure is the assumption that core heat flow statistics derived from mantle GCMs over the past 200 Ma are applicable at earlier times, as far back as the ICN. In addition, we are assuming that the small change in core temperature over this time interval does not affect either the dynamics of the lower mantle or the heat transfer through the mantle, thereby allowing us to use a fixed temperature CMB boundary condition for the mantle GCMs.

To test the validity of these assumptions, we show in Figure 6 the variation of CMB temperature and inner core radius versus age for Q_{cmb} =12 and 14 TW and zero radioactivity, Q_{rad} =0, calculated from the core evolution model described in the Supplementary Materials section using the parameters in Table 2. For these cases the decrease in the CMB temperature T_{cmb} since ICN is approximately 94^{o}K and the ICN age is 770 and 660 Ma, respectively. Figure 6 also shows the core evolution driven by the CMB heat flow from mantle GCM case 2 in Figure 1 reflected at 200 Ma then repeated periodically back in time, with 1 TW of heating from potassium-40 added to the outer core. This combination of thermal forcing increases the ICN age to 800 Ma. For these heat flows the core evolution model predicts outer core convective velocities of the order 10^{-3} m.s^{-1}, corresponding to convective overturn times of a few centuries. Clearly, the dynamic response time of the core is negligible compared to ICN age, and the decrease in CMB temperature since ICN is only 2%, a negligible amount in terms of its effect on the mantle GCM.

5 Inner core nucleation age

Figure 7 shows predicted ICN ages as functions of Q_{cmb} and melting curve parameter Θ for assumed values of present-day core radioactive heat production Q_{rad} of 0, 1, and 2 TW. In these calculations, the decay rate of radioactive potassium-40 was used. The boxes with dashed outlines delineate the parameter combinations that are allowed on the basis of our mantle GCM heat flow statistics and melting relations for inner core compositions (Anzellini et al., 2013) The lower limit of the dashed boxes correspond to a Grüneisen parameter of γ=0.9 and the upper limits corresponds to γ=1.8. The dotted lines indicate the 0-200 Ma mean CMB heat flow from our mantle GCMs.

Without radioactive heating, ICN ages range from more than 1600 Ma for $Q_{cmb} = 6$ TW to less than 400 Ma for $Q_{cmb} = 18$ TW (Figures 7a,b), but using just the allowed values of Q_{cmb} and Θ limits this range to 400-950 Ma. As the present-day radioactive heat content increases, the predicted age of ICN also increases, but the change is rather small for the amounts of radioactive heating that are probable in the core. High-pressure partition experiments indicate solubility of potassium in core alloys (Bouhifd et al., 2007) but the upper limit on its heat production in the core appear to be substantially less than 1 TW (Hirose et al., 2013; Watanabe et al., 2014). Similarly, high-pressure partition experiments on uranium (Malavergne et al., 2007) indicate that its maximum heat production in the core is also substantially less than one terawatt. Therefore, taking 1 TW as an upper bound on total radioactive heat production in the core, the maximum ICN age within the dashed boxes in Figure 7c is about 1100 Ma.

There is an additional constraint on core evolution related to its ability to sustain the geodynamo, which further restricts inner core age. Since we know that the geomagnetic field has persisted since 3400 Ma at least (Tarduno et al., 2010) the energetics of the core must allow for dynamo action today, just after ICN, as well as before ICN. The shaded regions in Figure 7 denote parameter combinations for which the core is subcritical for convection-driven dynamo action today (unshaded), supercritical for dynamo action today (yellow), supercritical for dynamo action 50 Myr after ICN (brown), and supercritical for dynamo action just prior to ICN (red). These regions are defined in terms of a prediction of the magnetic Reynolds number of convection in the outer core based on scaling laws derived from the systematics of numerical dynamos (Christensen and Aubert, 2006). Here the predicted magnetic Reynolds number of the outer core Rm is calculated using a method developed by Aubert et al. (2009) in which

$$Rm \simeq 1.31 p^{0.42} Pm \quad (11)$$

where p is the (dimensionless) power from convection available to drive the dynamo. The relationship between p and core parameters is given in the Supplementary Material. The critical value for dynamo action in a fully convective outer core is $Rm_{crit} \geq 40$ (Christensen et al., 1999); the criterion based on (11) used for the shadings in Figure 7 is Rm=100.

The boundaries separating subcritical and supercritical dynamo regimes depend sensitively on the thermal conductivity of the core because the adiabatic heat flux, which controls thermal convection in the outer core, is proportional to thermal conductivity. The buoyancy flux at the CMB is thermal and depends on the global mean heat flux relative to the heat flux down the adiabatic gradient there. Accordingly, if core thermal conductivity is high, the average CMB heat flux in the core is subadiabatic and makes a negative contribution to convective power p. Strongly subadiabatic CMB conditions reduce p to the point where $Rm < Rm_{crit}$, indicating dynamo failure. Furthermore, a key assumption used to derive (11), that the outer core is adiabatic (well-mixed) outside of boundary layers, is no longer valid in the strongly stratified regime, casting further doubt on the viability of such a convective dynamo.

In Figure 7a,b, two thermal conductivities are considered, k= 100 and 130 W.m^{-1}.K^{-1}. The lower value is representative of the core conductivity predicted by Zhang et al. (2015) on the basis of density functional theory (DFT) including electron-electron scattering; the higher value is more representative of DFT calculations without this effect (Pozzo et al., 2014). The left hand portion of every panel has $Rm < Rm_{crit}$, implying that, for the oldest inner core ages, the present-day core would be incapable of sustaining the geomagnetic field by thermochemical convection. The situation improves moving to the right Figure 7, where the present-day core is supercritical for convective dynamo action for most parameter combinations. Problems for the geodynamo reappear, however, when considering the state of the core shortly after and before ICN. The darkest (red) shadings in Figure 7 indicate (Q_{cmb},Θ) combinations for which the core is supercritical for convective dynamo action just prior to ICN. This region includes only large Q_{cmb}-values and generally young inner core ages. Figures 7a,b show that the maximum inner core age for which the geodynamo would be supercritical prior to ICN with Q_{rad}= 0 are approximately 775 Ma for k= 100 W.m^{-1}.K^{-1}, and for this $Q_{cmb} \geq 12$ TW is needed. For k= 130 W.m^{-1}.K^{-1}, the maximum IC age is only about 550 Ma, and in this case $Q_{cmb} \geq$16 TW is needed before ICN. Figure 7c indicates the maximum IC age increases by only 80 Ma with Q_{rad}=1 TW.

To further demonstrate this point, we show in Figure 7d the ICN ages predicted for Q_{rad}=2 TW. Although this amount of radioactive heating is not supported by partition experiments or by cosmochemical considerations (McDonough, 2003) it is nevertheless of some theoretical interest because whole-Earth thermal history calculations reveal that the increase in heat production with age corresponding to this amount of potassium in the present-day core helps the geodynamo survive back to 3.4 Ga (Driscoll and Bercovici, 2014). Nevertheless, it would increase the allowable IC age by only about 160 Ma, strengthening our inference of a young inner core. Unless the amount of core radioactive heating greatly exceeds current estimates, the ICN was a relatively recent event; within 800 Ma if there is no radioactive heating in the core, and within 1100 Ma, even if radioactive heating is abundant. By the same token, our models permit inner core ages as young as 400 Ma.

6 Implications for powering the geodynamo and inner core convection

The combination of our mantle GCMs and the k=100 W.m^{-1}.K^{-1}core evolution cases in Figure 7 provides a self-consistent (although non-unique) picture of core-mantle thermal interaction from the present-day backward in time to the ICN. With this combination, our mantle GCMs predict supercritical convective dynamo conditions at the present-day, just after ICN, and also just before ICN, although with much reduced power. In contrast, according to Figure 7b, our mantle GCMs do not provide enough heat flow to power the geodynamo by thermal convection prior to ICN if k=130 W.m^{-1}.K^{-1}. It is possible that CMB heat flow was larger before ICN compared to 0-200 Ma, but it seems coincidental that CMB heat flow would change appreciably just at the time of ICN. Another possibility is that CMB heat flow today is actually a lot larger than our mantle GCMs predict. Apart from implying a very young inner core – a Paleozoic or possibly Mesozoic ICN– the consequences of this situation have hardly been explored.

The results in Figure 7 also bear on the question of subsolidus thermal convection within the inner core, which depends on whether the temperature profile in the inner core is subadiabatic or superadiabatic. The thermal state of the inner core is governed by a competition between cooling at the ICB and diffusion of the inner core internal heat, with fast inner core growth and low thermal conductivity leading to steeper and hence less stable temperature profiles. Deguen et al. (2011) showed that the inner core temperature profile is expected to be superadiabatic if

$$\frac{dr_{icb}^2}{dt} > 6\kappa_{\rm ic}\left(\frac{dT_{melt}}{dT_{ad}} - 1\right)^{-1}, \tag{12}$$

where dT_{melt}/dT_{ad} is the ratio of the Clapeyron slope dT_{melt}/dP over the adiabatic gradient dT_{ad}/dP, and $\kappa_{\rm ic}$ is the thermal diffusivity in the inner core. If the inner core is assumed to grow as $r_{icb} \propto \sqrt{t}$ (Labrosse, 2014), a reasonable approximation to the growth curves in Figure 6, then (12) can be re-written as a criterion on the maximum ICN age τ_{ICN} that would generate a superadiabatic temperature profile in the inner core:

$$\tau_{ICN} < \frac{r_{icb}^2}{6\kappa_{\rm ic}}\left(\frac{dT_{melt}}{dT_{ad}} - 1\right) \tag{13}$$

(Deguen et al., 2011).

The thermal conductivity in solid iron at inner core conditions is likely to be even larger than for liquid iron at CMB conditions, with some estimates exceeding 170 W.m^{-1}.K^{-1} (de Koker et al., 2012; Pozzo et al., 2012; Gomi et al., 2013; Pozzo et al., 2014), which corresponds to $\kappa_{\rm ic} > 1.7\,10^{-5}$ m.s^{-2}. Assuming this conductivity and using $dT_{melt}/dT_{ad} \simeq 1.6$, (13) gives the maximum ICN age for inner core superadiabaticity of $\tau_{ICN} \leq 270$ Ma. As this maximum is smaller than our most extreme ICN age estimates, such high thermal conductivity implies that the inner core is thermally stably stratified and therefore subsolidus thermal convection in the inner core would be unlikely. In contrast, the lower conductivity value of k=100 W.m^{-1}.K^{-1} recently obtained by Zhang et al. (2015) leads to a different interpretation. With this lower conductivity the critical ICN age for subsolidus convection in the inner core increases to $\simeq$ 460 Ma. Given the range of ICN ages our mantle GCMs predict (400-1100 Ma), convection in the inner core becomes marginally possible.

7 Implications for mantle circulation, past and future

The core evolution calculations in the previous sections could be extended to greater age, however it would be necessary to couple the core evolution more directly to the mantle evolution, allowing the CMB temperature to change with time, and in addition, assumptions would be needed regarding the surface tectonic boundary conditions and the possibility of mantle melting. Because it is not possible to reconstruct global plate distributions in the deep past and our mantle GCMs do not include melting, we have restricted our attention to the time since ICN. Coupled mantle-core thermal evolution calculations that do not make use of plate motions but do include time dependent mantle convection and dynamo thermodynamics (Nakagawa and Tackley 2013, 2014) generally come to the same conclusions as we have regarding the time of ICN.

Not only has the geomagnetic field persisted for 3.4 Ga at least (Tarduno, 2010), there is no paleomagnetic evidence that the geodynamo ever shut off (Biggen et al., 2012). Assuming k=100 W.m^{-1}.K^{-1} (Zhang et al., 2015), the time average core heat flow in our mantle GCMs is adequate to maintain convective dynamo conditions from the present-day to some time before inner core nucleation, although slightly more core heat flow would be needed for thermal convection in the deep past when the core temperature, the adiabatic gradient, and the rotation rate were higher. Although the plate tectonics Wilson cycle may only date back to 3 Ga (Shirey and Richardson, 2011) the greater antiquity of the geodynamo implies that some form of global mantle circulation was operational before then, extracting heat from the core at rates comparable to or larger than the past 200 Ma. As for the future, our models predict that, at the present rate of heat loss to the mantle, a large part of the outer core will remain molten for more than one Gyr and supercritical convective dynamo conditions will prevail over that time.

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Supplementary Material

The evolution of the inner core radius can be written as

$$\dot{r}_{icb} = \frac{Q_{cmb} - Q_{rad}}{P}. \tag{14}$$

with $P = P_l + P_g + P_s$ is the sum of contributions from latent heat release, gravitational energy release and release of sensible heat. Individually these can be expressed as (Labrosse; 2003)

$$P_l = 4\pi r_{\text{icb}}^2 \rho(r_{\text{icb}}) T_{melt}(r_{\text{icb}}) \Delta S, \tag{15}$$

$$P_g = \frac{8\pi^2}{3} \mathcal{G} \Delta\rho \, \rho_c r_{\text{icb}}^2 r_{\text{cmb}}^2 \left(\frac{3}{5} - \frac{r_{\text{icb}}^2}{r_{\text{cmb}}^2} \right), \tag{16}$$

$$P_s = 4\pi H^3 \rho_c c_p T_{meltc} \left(1 - \frac{2}{3\gamma} \right) \frac{r_{\text{icb}}}{r_T^2} \exp\left[\left(\frac{2}{3\gamma} - 1 \right) \frac{r_{\text{icb}}^2}{r_T^2} \right] I(H, r_{\text{cmb}}), \tag{17}$$

where the radial profiles of density ρ, gravity g, melting temperature T_{melt}, and temperature $T = T_{ad}$ in the outer core are given by

$$\rho = \rho_c \exp\left(-\frac{r^2}{r_\rho^2} \right), \tag{18}$$

$$g = \frac{4\pi}{3} \mathcal{G} \rho_c r \left(1 - \frac{3}{5} \frac{r^2}{r_\rho^2} \right), \tag{19}$$

$$T_{melt} = T_{meltc} \exp\left[-2 \left(1 - \frac{1}{3\gamma} \right) \frac{r^2}{r_T^2} \right], \tag{20}$$

$$T = T_{melt}(r_{icb}) \exp\left(\frac{r_{icb}^2 - r^2}{r_T^2} \right), \tag{21}$$

with

$$r_\rho = \sqrt{\frac{3K_0}{2\pi \mathcal{G} \rho_0 \rho_c} \left(\ln \frac{\rho_c}{\rho_0} + 1 \right)}, \quad r_T = \sqrt{\frac{3c_p}{2\pi \alpha_c \rho_c \mathcal{G}}}. \tag{22}$$

Here T_{meltc} is the melting temperature at the center of the core, r_{icb} the radius of the inner core, γ the Grüneisen coefficient assumed constant, ρ_0 and ρ_c are the density of liquid core material at zero pressure and at the center of the core, respectively, K_0 the incompressibility at zero pressure, $\mathcal{G}$ the gravitational constant, c_p the heat capacity assumed constant, and α_c the coefficient of thermal expansion of liquid core material at the center of the core, ΔS is the entropy of melting, r_{cmb} is the CMB radius, $\Delta\rho$ is the density difference between inner and outer core due to differences in their light element contents, $H = \left(1/r_\rho^2 + 1/r_T^2\right)^{-1/2}$, and

$$I(H, r_{\text{cmb}}) = \frac{\sqrt{\pi}}{2} \text{erf}\left(\frac{r_{\text{cmb}}}{H} \right) - \frac{r_{\text{cmb}}}{H} \exp\left(-\frac{r_{\text{cmb}}^2}{H^2} \right). \tag{23}$$

Integrating (14) backward in time from present-day conditions using the parameters in Table 2 with assumed values of Q_{cmb} and Q_{rad} gives $r_{\rm icb}(t)$. If χ represents light element concentration in the well-mixed outer core, evolution of the average light element concentration with complete partitioning (that is assuming no light elements partition into the inner core) can be approximated by (Olson et al., 2013)

$$\dot{\chi}(t) \simeq 3\chi \frac{r_{\rm icb}^2 \dot{r}_{\rm icb}}{r_{\rm cmb}^3} \tag{24}$$

which completes the core evolution model.

For the dynamo, the various buoyancy sources in the outer core are defined by the following five Rayleigh numbers:

$$Ra = \frac{\beta g D^5 \dot{\chi}}{\nu^2 \kappa}, \quad Ra_q = \frac{\alpha g D^4 (\bar{q}_{\rm cmb} - q_{\rm ad})}{\nu \kappa k}, \tag{25}$$

and

$$Ra_{q'} = \frac{\alpha g D^4 \delta q_{cmb}}{\nu \kappa k}, \quad Ra_h = \frac{\alpha g D^3 h}{c_p \nu^2 \kappa}, \quad Ra_\theta = -\frac{\alpha g d^5 \dot{\theta}_{ad}}{\nu^2 \kappa}. \tag{26}$$

where h is the volumetric heat source density and $\dot{\theta}_{ad} = \frac{\partial T_{\rm ad}}{\partial t} - \kappa \nabla^2 T_{\rm ad}$. In terms of these, the sink term in the codensity transport equation can be written approximately as

$$\epsilon \simeq -1 + \frac{Ra_\theta}{Ra} + \frac{Ra_h}{Ra} \simeq -1.5. \tag{27}$$

The global mean heat flux boundary condition at the CMB in terms of the dimensionless codensity is given by

$$\left.\frac{\partial C^*}{\partial r^*}\right|_{\rm cmb} = -\frac{Ra_q}{Ra}. \tag{28}$$

We use the following scaling law for the magnetic Reynolds number of the convection in the outer core, derived from numerical dynamos (Aubert et al., 2009).

$$Rm \simeq 1.31 p^{0.42} Pm \tag{29}$$

where p is the (dimensionless) power from convection available to drive the dynamo. Aubert et al. (2009) related p to a modified Rayleigh number defined as

$$Ra_Q = \frac{g_o F}{4\pi \rho \Omega^3 D^4} \tag{30}$$

in which F is the sum of the buoyancy productions at the ICB and CMB, according to

$$p \simeq (cRa_Q)^{0.42}. \qquad (31)$$

Here the factor c is meant to absorb the effects of inner core size and stratification. For dynamos with destabilizing buoyancy fluxes at both boundaries, and also for dynamos with slightly stabilizing buoyancy flux at the CMB, Aubert et al. (2009) find that c increases as r_i decreases, with $c \simeq 0.4$ for the present-day inner core size and $c \simeq 1$ near inner core nucleation. We use these values in estimating the magnetic Reynolds number Rm. In our notation, Ra_Q is related to Ra and the boundary conditions on codensity according to.

$$Ra_Q = -Pr^{-1}E^2 Ra\left(r_{cmb}^{*2} \left.\frac{\partial C^*}{\partial r^*}\right|_{\mathrm{cmb}} + r_{icb}^{*2} \left.\frac{\partial C^*}{\partial r^*}\right|_{\mathrm{icb}} \right). \qquad (32)$$

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Table 1: Mantle GCM Parameters

Parameter	Notation	Value
Superadiabatic temperature difference	ΔT_m	2500 K [a]
Reference viscosities: plate, upper mantle, lower mantle	$\eta_{p,u,l}$	$1200, 0.6, 100 \times 10^{20}$ Pa.s [b]
Radioactive heat production	h_m	2.5×10^{-8} $W.m^{-3}$ [c]
Reference density	ρ_m	3300 kg.m^{-3}
Initial D" dense layer thickness	d_0	250 km [d]
Initial D" density anomaly	$\Delta\rho_{d0}$	82.5 kg.m^{-3} [c]
Heat capacity	c_m	1000 J.kg^{-1}.K^{-1} [f]
Thermal expansion coefficient	α_m	2×10^{-5} K^{-1} [f]
Thermal conductivity above the CMB	k_m	4 W.m^{-1}.K^{-1} [c]
Surface radius	$r_{\rm surf}$	6371 km
CMB radius	$r_{\rm cmb}$	3480 km
Viscosity activation energy	E	190 kJ.mol^{-1} [e]

[a] Boehler et al., (1995); [b] Simons and Hager (1997); [c] Zhang et al. (2010); [d] Wang and Wen (2004); [e] van Hunen et al. (2005); [f] Schubert et al. (2001).

Table 2: Core Evolution Parameters

Parameter	Notation	Value [*Present-day]
Density at core center	ρ_c	12500 kg.m^{-3} [a]
Density at zero pressure	ρ_0	7500 kg.m^{-3}
Compositional density jump at the ICB	$\Delta\rho$	500 kg.m^{-3}* [b]
Incompressibility at zero pressure	K_0	4.75×10^{11} Pa
Melting temperature at the ICB	T_{melt}	5500 K* [c]
Entropy of melting	ΔS	120 J.kg^{-1}.K^{-1} [d]
Grüneisen parameter	γ	1.5 [e]
Heat capacity	c_c	850 J.kg^{-1}.K^{-1} [e]
Thermal expansion coefficient	α_c	1.3×10^{-5} K^{-1} [e]
Compositional expansion coefficient	β	1
Thermal conductivity at the CMB	k	100, 130 W.m^{-1}.K^{-1}
Density length scale	r_ρ	7400 km[a]
Temperature length scale	r_T	6040 km [c]
ICB radius	r_{icb}	1221 km*[a]
CMB radius	r_{cmb}	3480 km[a]
Outer core light elements	χ	9.8 wt.%* [f]
Outer core kinematic viscosity	ν	10^{-6} m^2.s^{-1} [g]

[a] Dziewonski and Anderson (1981); [b] Masters and Gubbins (2003); [c] Ancellini et al. (2013); [d] Poirier (1990); [e] Vocadlo et al. (2003); [f] Hirose et al. (2013); [g] Perriallt et al. (2010).

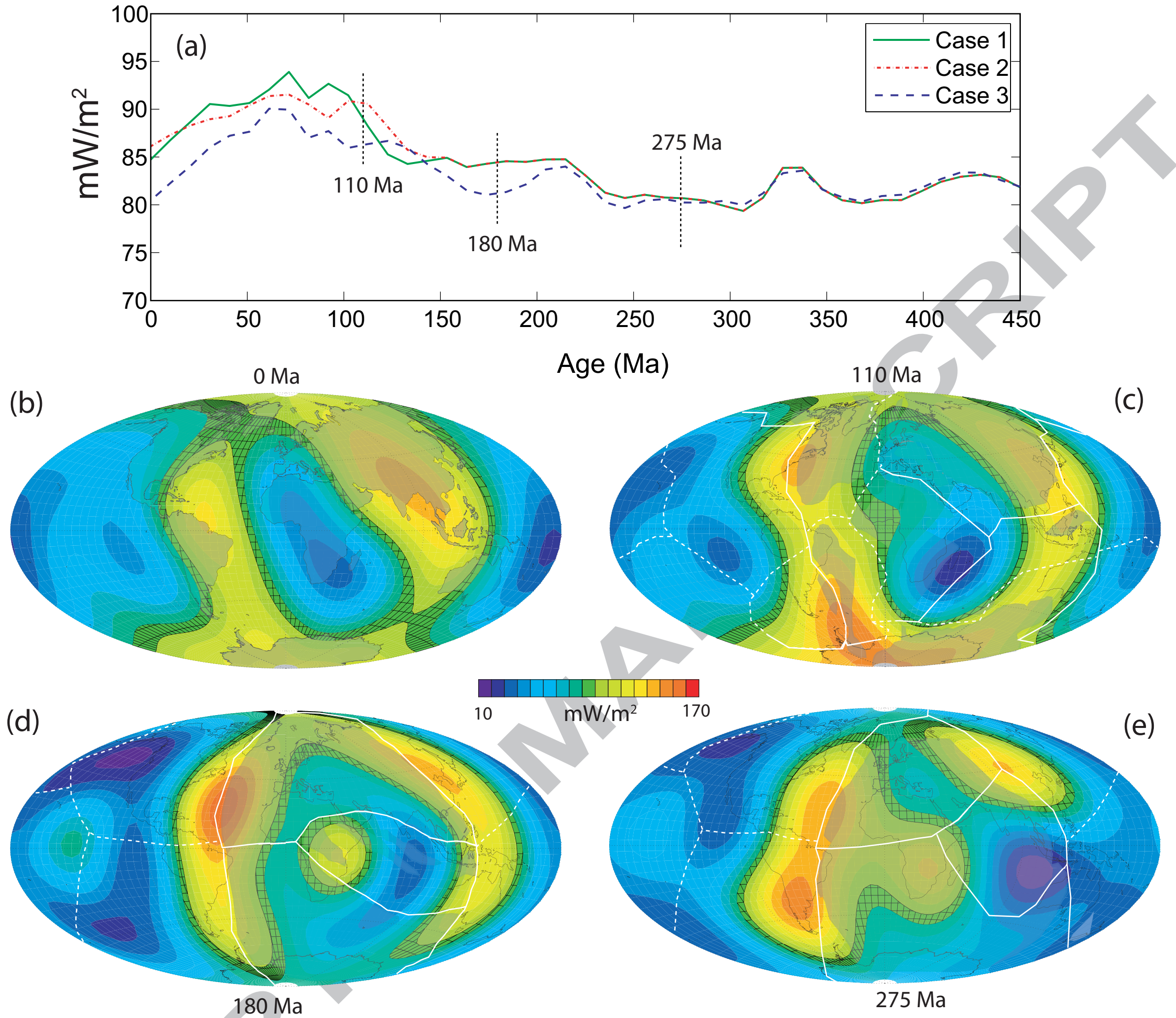

Fig 1

Figure 1. Heat flux on the core-mantle boundary (CMB) from mantle global circulation models (GCMs). (a): Time series of global mean CMB heat flux versus age from three mantle GCMs (Rudolph and Zhong, 2014) using plate reconstructions by Muller et al. (2008; 0-140 Ma; Case 1), Lithgow-Bertelloni and Richards (1998; 0-119 Ma; Case 2), and Seton et al., (2012; 0-200 Ma; Case 3). (b)-(e): Snapshots of CMB heat flux patterns for the present-day and the three Case 2 epochs labeled on the time series. The hashed contours enclose regions with CMB heat flux of 100 mW.m^{-2} or more. Continents (shaded) and reconstructed plate boundaries (solid=convergent; dashed=divergent) are shown for reference.

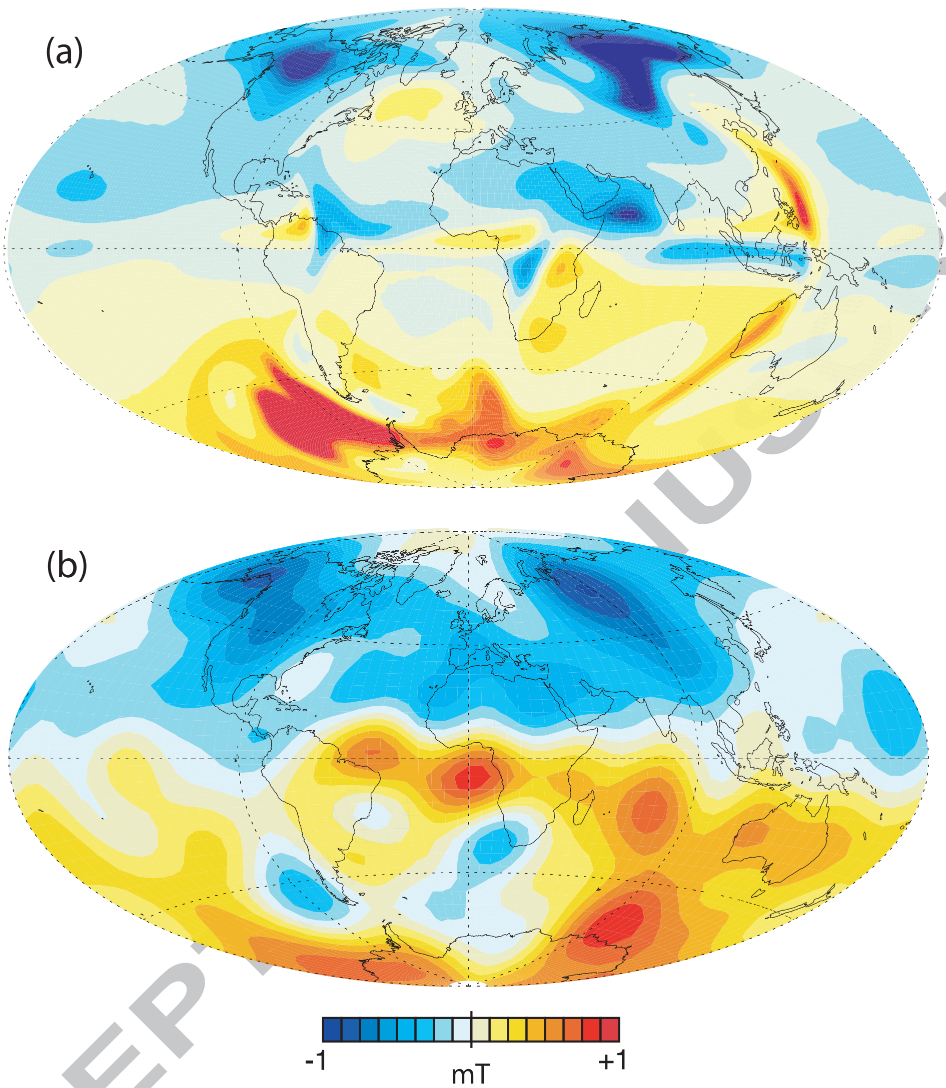

Fig 2

Figure 2. Comparison between (a) a snapshot of the radial magnetic field on the CMB from a numerical dynamo driven by the present-day 0 Ma pattern of CMB heat flux shown in Figure 1b and (b) the present-day geomagnetic field intensity on the CMB in millitesla (mT) from core field model POMME 2008 (http://geomag.org/index.html). Dynamo magnetic field intensity is in dimensionless Elsasser units defined in the text.

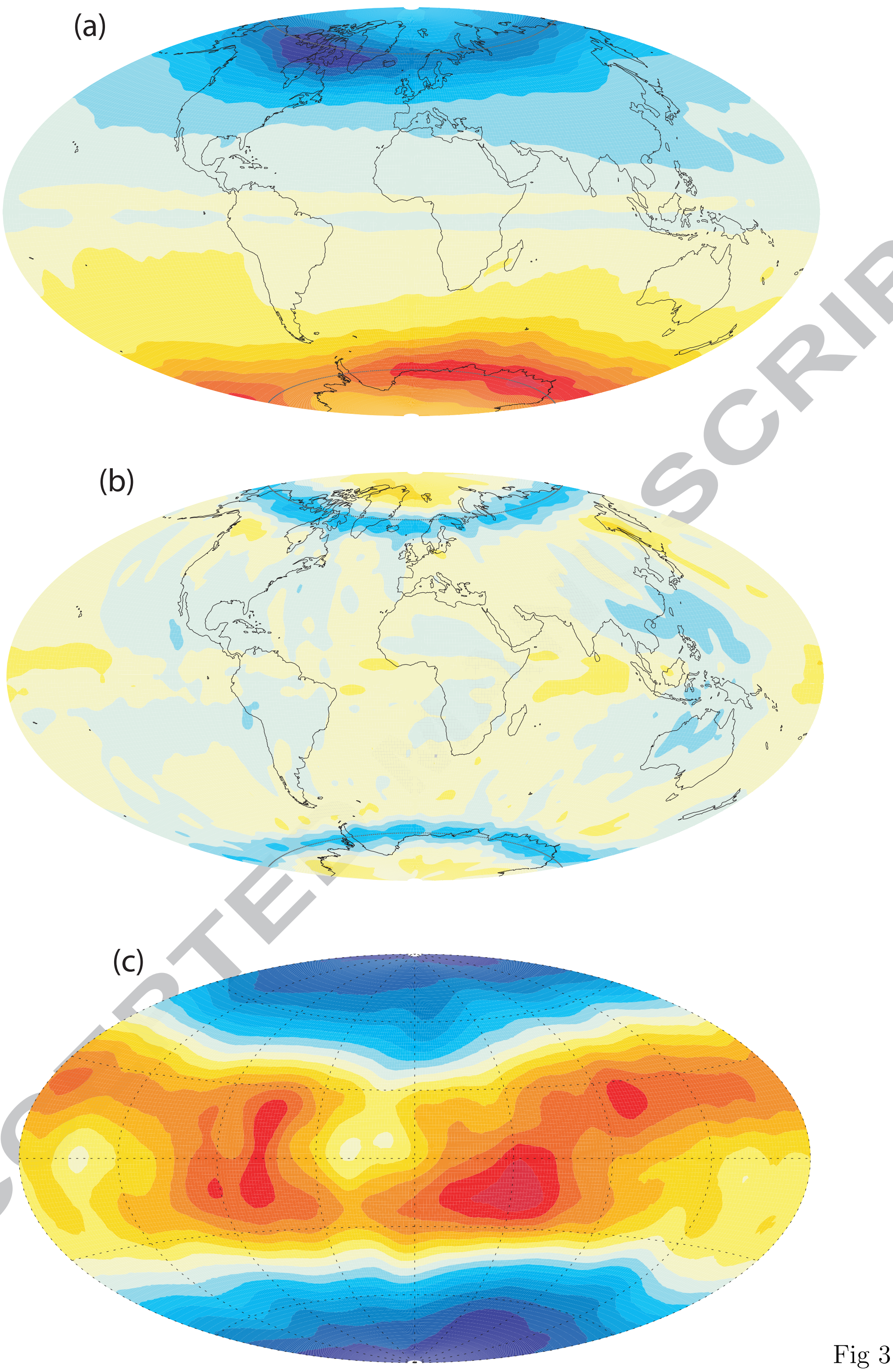

Figure 3. Time average structure of the numerical dynamo in Figure 2. (a) Time average radial magnetic field on the CMB; contours in 0.2 dimensionless units. (b) Time average radial fluid velocity at a distance $z = 0.05D$ below the CMB, contours in magnetic Reynolds number units of 3. (c) Time average codensity flux on the ICB, contoured in 0.1 dimensionless units, oriented with maps (a) and (b).

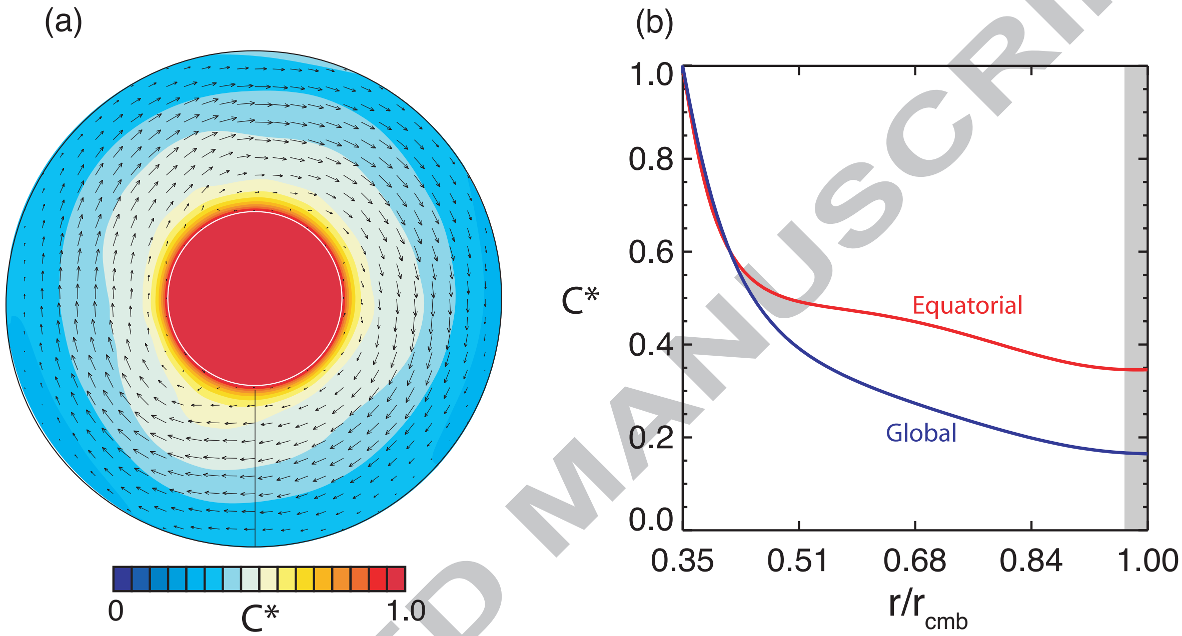

Fig 4

Figure 4. (a) Time average of the dimensionless codensity in the equatorial plane of the dynamo shown in Figures 2 and 3 with time average velocity arrows superimposed. Thin line marks 0^o longitude. (b) Global and equatorial averages of the radial variation of codensity from the same numerical dynamo, including the thin shaded region with a slightly stable stratification beneath the CMB in the equatorial average.

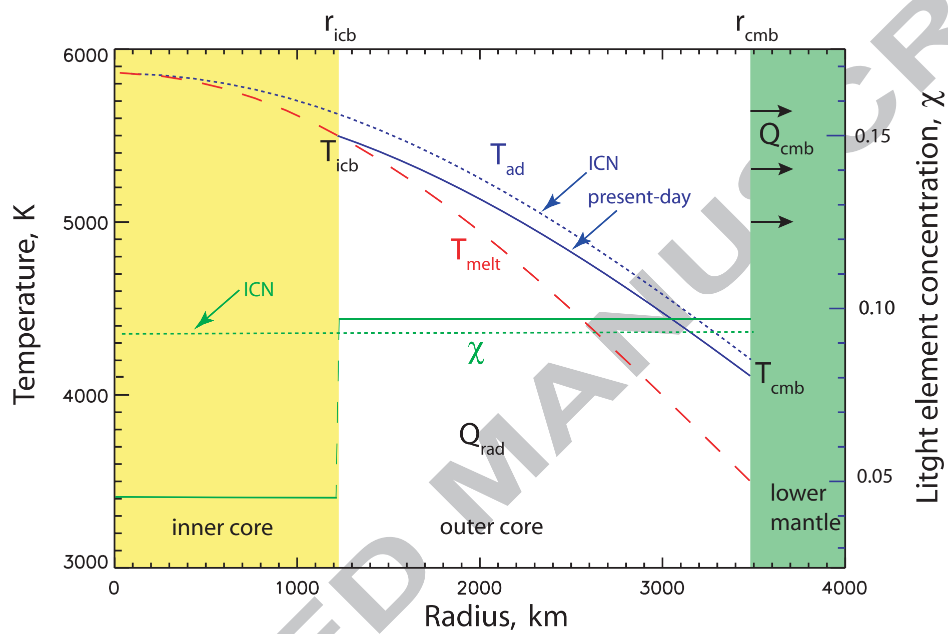

Fig 5

Figure 5. Evolution model of the core. Solid curves show present-day profiles of adiabatic temperature T_{ad} and light element concentration χ; dotted curves show these profiles at the time of inner core nucleation, ICN. Dashed curve T_{melt} is a representative melting curve in the core. Q_{cmb} and Q_{rad} are total CMB heat flow and internal radioactive heat production, respectively.

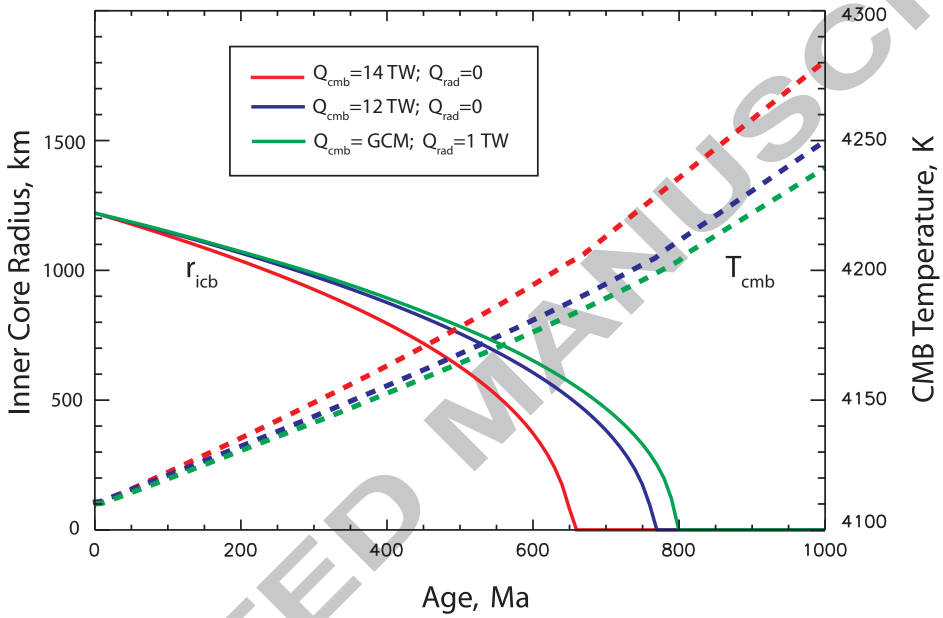

Fig 6

Figure 6. Evolution of the core for different values of the total CMB heat flow, assumed constant in time. T_{cmb} and r_{icb} denote CMB temperature and inner core radius, respectively.

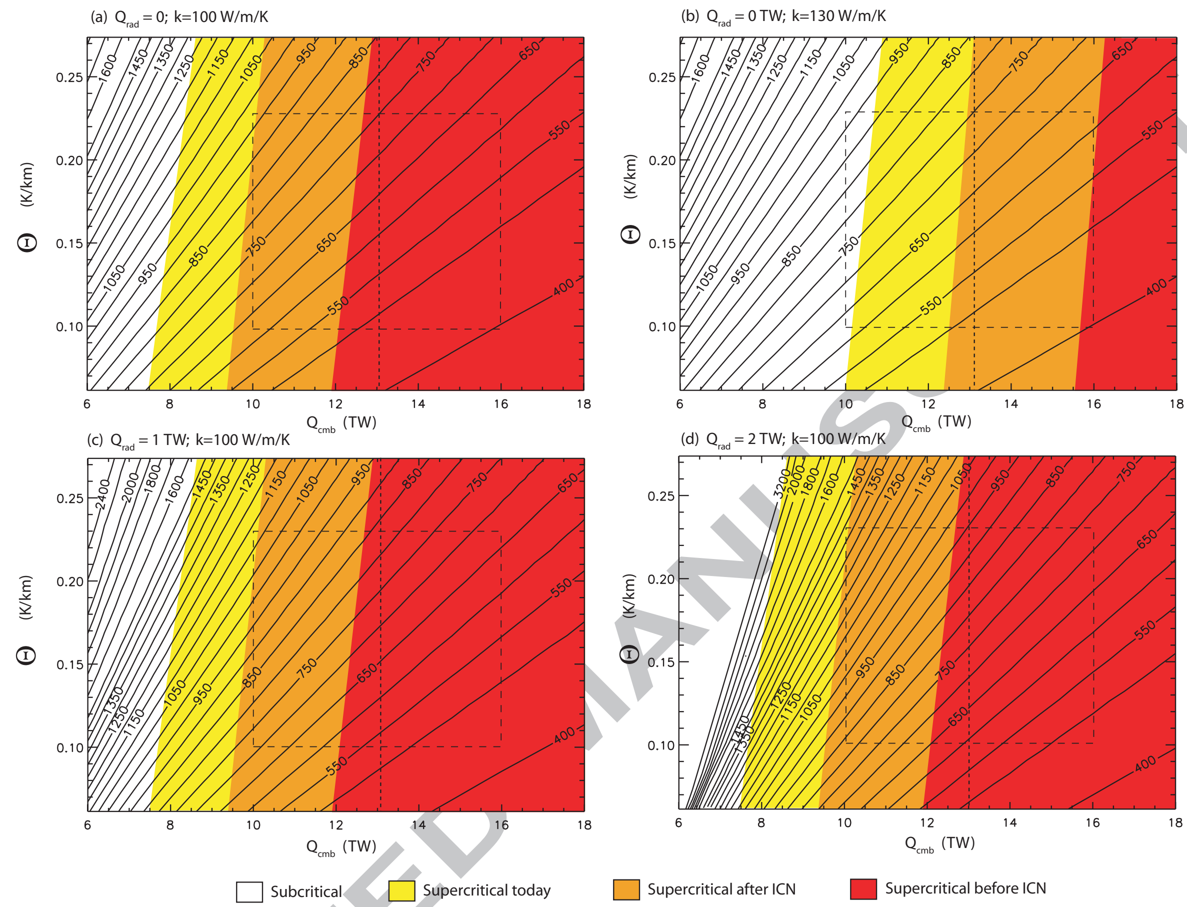

Fig 7

Figure 7. Predicted ages of inner core nucleation (ICN) in millions of years as a function of total CMB heat flow Q_{cmb} and Θ, the difference between the slope of the melting curve and the adiabat at the inner core boundary, calculated using different combinations of present-day potassium-40 radioactive heat production Q_{rad} and outer core thermal conductivity k. Panels a, c, and d use k=100 W.m^{-1}.K^{-1}; Panel b uses k=130 W.m^{-1}.K^{-1}; Top row (a & b) use Q_{rad}=0; Bottom row: (c,d) use Q_{rad}=(1,2) TW, respectively. Shadings correspond to dynamo states: white=subcritical; yellow=supercritical today; light brown=supercritical 50 Myr after ICN; red=supercritical just prior to ICN. Dashed boxes indicate allowed region based on mantle GCMs and core melting relations. Dotted lines indicate the time average Q_{cmb} from our mantle GCMs.