

HD 191939: Three Sub-Neptunes Transiting a Sun-like Star Only 54 pc Away

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ABSTRACT

We present the discovery of three sub-Neptune-sized planets transiting the nearby and bright Sun-like star HD 191939 (TIC 269701147, TOI 1339), a $K_s = 7.18$ magnitude G8V dwarf at a distance of only 54 parsecs. This multi-planetary system is one of the best candidates for atmospheric characterization to date, with all its known planets suitable for multi-wavelength transmission spectroscopy. We validate the planetary nature of the transit signals by combining five months of data from the *Transiting Exoplanet Survey Satellite* with follow-up ground-based photometry, archival optical images, radial velocities, and high angular resolution observations. The three sub-Neptunes have similar radii ($R_b = 3.37^{+0.13}_{-0.13} R_\oplus$, $R_c = 3.22^{+0.15}_{-0.14} R_\oplus$, and $R_d = 3.16^{+0.11}_{-0.11} R_\oplus$) and their orbits are consistent with a stable, circular, and co-planar architecture near mean motion resonances of 1:3 and 3:4 ($P_b = 8.88$ days, $P_c = 28.58$ days, and $P_d = 38.35$ days). The HD 191939 system is an excellent candidate for precise mass determinations of the planets with high-resolution spectroscopy due to the host star’s brightness and low chromospheric activity. Moreover, the system’s compact and near-resonant nature can provide an independent way to measure planetary masses via transit timing variations while also enabling dynamical and evolutionary studies. Finally, as a prime target for transmission spectroscopy of all three planets’ atmospheres, HD 191939 can offer a unique insight into multiple sub-Neptunes born from a proto-planetary disk that may have resembled that of the early Sun.

Keywords: Planetary systems, planets and satellites: detection – stars: individual (HD 191939, TIC 269701147, TOI 1339) – techniques: transit photometry

1. INTRODUCTION

The *Transiting Exoplanet Survey Satellite* (*TESS*, Ricker et al. 2014) was designed to detect transiting “super-Earths” ($R_p = 1.25 - 2 R_\oplus$, $M_p \approx 1 - 10 M_\oplus$) and “sub-Neptunes” ($R_p = 2 - 4 R_\oplus$, $M_p \approx 10 - 40 M_\oplus$) around the nearest and brightest main-sequence stars. As a result, planets detected by *TESS* will be some of the best candidates for follow-up spectroscopy and future atmospheric characterization studies.

Since the beginning of science operations in 2018, *TESS* has discovered about a dozen multi-transiting planet systems, (e.g. Huang et al. 2018; Dragomir et al. 2019; Günther et al. 2019; Quinn et al. 2019), including some of the brightest known to date, thus yielding prime targets for detailed characterization (e.g. Huang et al. 2018; Dragomir et al. 2019; Günther et al. 2019; Quinn et al. 2019). These “multis” offer a valuable opportunity to perform comparative exoplanetology and learn about the planetary formation and evolutionary processes in the controlled environment of their host star. In addition, they often have a greater scientific potential than single-planet systems because they can be characterized comprehensively by exploiting the power of multiple techniques (Ragozzine & Holman 2010). For instance, transit photometry provides the planets’ sizes and orbital periods; radial velocity (RV) observations and transit timing variations (TTV) constrain planetary masses and orbital architectures; and transmission spectroscopy enables studies of planetary atmospheres, compositions, and habitability regimes.

The *Kepler* mission revealed that multi-transiting planet systems are ubiquitous (Latham et al. 2011; Lissauer et al. 2011; Lissauer et al. 2014; Rowe et al. 2014), particularly in the super-Earth to mini-Neptune regime (e.g. Howard et al. 2010; Fressin et al. 2013). Despite their widespread occurrence, the majority of *Kepler* multis are too faint, distant, and/or small to precisely determine planetary masses with independent RV surveys. Consequently, many of them lack mass and density measurements. With *TESS*, however, the population of multis amenable to follow-up studies will grow as nearer and brighter systems are detected. Many of these *TESS* discoveries will be sub-Jovian-sized planets well suited for spectroscopic studies of planetary masses (e.g. Cloutier et al. 2018) and atmospheres (e.g. Kempton et al. 2018) due to their larger sizes and their host star’s proximity and brightness.

Here we focus on HD 191939 (TOI 1339, TIC 269701147), a bright ($V = 8.97$ mag, $K = 7.18$ mag), nearby ($d = 53.69 \pm 0.07$ pc), Sun-like (G8V) star with a radius of $R_* = 0.945 \pm 0.021 R_\odot$, a mass of $M_* = 0.92 \pm 0.06 M_\odot$, and a temperature of $T_{\text{eff}} = 5400 \pm 50$ K. Using *TESS* data from sectors 15–19, we present the discovery of three sub-Neptune-sized planets around HD 191939 and validate their transit signals with archival optical images, RVs, ground-based photometric follow-up, and high-resolution imaging. At a distance of only 54 pc, HD 191939 is one of the nearest and brightest multi-transiting planet systems known to date (see Figure 1). Due to the host star’s proximity, brightness, and low chromospheric activity, this multi

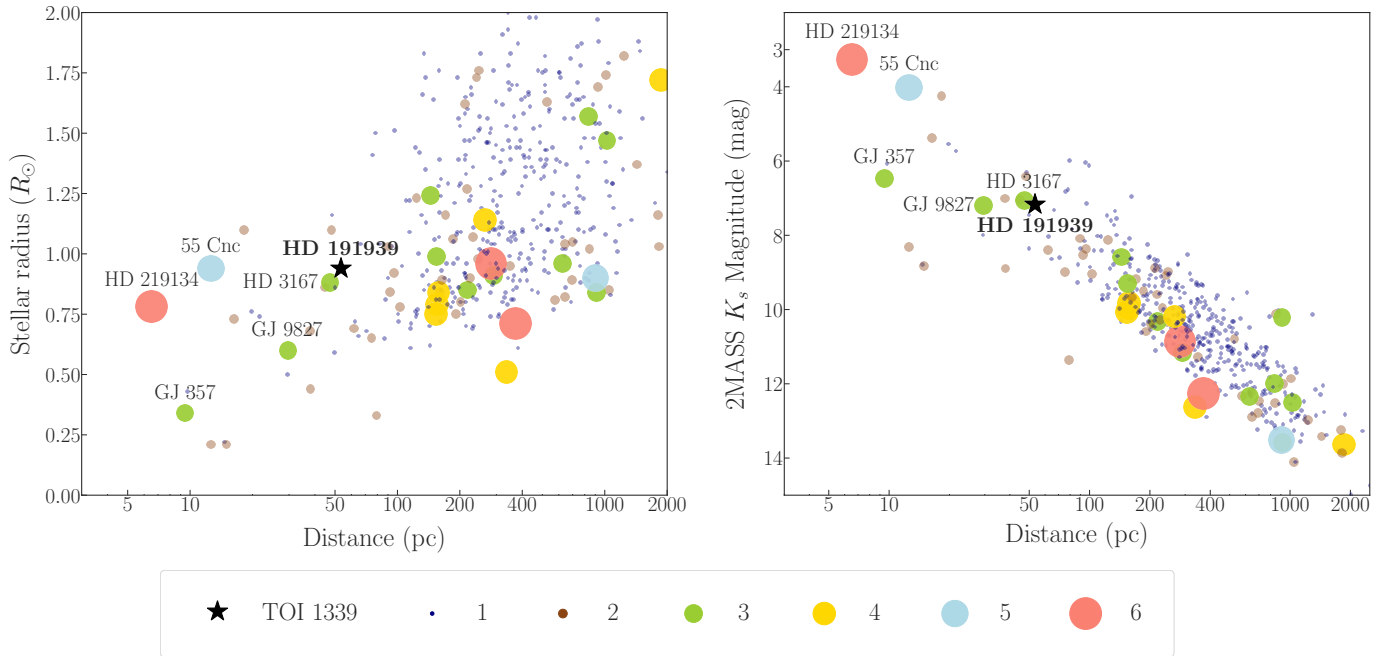


Figure 1. Host star radius and K_s magnitude of confirmed single- and multi-planetary systems. In these views, we only show systems with measured masses and a relative error in planet mass, planet radius, and host star radius less than 30%. Names are only displayed for stars with at least three transiting planets and with distances less than 100 pc. The size and color of each system depends on the number of planets that it hosts. Data were retrieved from the NASA Exoplanet Archive on 2020 February 1.

is a prime target for follow-up photometric and spectroscopic studies. As we step into the era of the *James Webb Space Telescope* (JWST), HD 191939 will be an excellent candidate for detailed atmospheric characterization as well.

This paper is organized as follows. Section 2 presents the *TESS* photometry and the available optical, photometric, and spectroscopic observations of HD 191939. In Section 3, we constrain the stellar parameters of HD 191939. Section 4 examines multiple false positive scenarios and confirms the planetary nature of the *TESS* transit signals. In Section 5, we describe our transit fitting routine, determine the system’s physical and orbital parameters, investigate its dynamical properties, and discuss its prospects for atmospheric characterization. Section 6 places HD 191939 in the context of known planetary systems and highlights possible research avenues to improve our current knowledge of HD 191939. Finally, we summarize our results and present our conclusions in Section 7.

2. OBSERVATIONS

2.1. *TESS* Photometry

With a *TESS* magnitude of $T = 8.29$ mag, a radius smaller than the Sun’s ($R_* = 0.945 \pm 0.021 R_\odot$), and a low contaminating ratio (~ 0.005), HD 191939 was included in the *TESS* Candidate Target List as

a high-priority target (Stassun et al. 2018). As such, HD 191939 was pre-selected for 2-min. observations consisting of 11×11 pixels sub-arrays centered on the target. The star’s astrometric and photometric properties are listed in Table 1.

The *TESS* spacecraft observed HD 191939 (RA J2015.5 = 20:08:06.150, Dec J2015.5 = +66:51:01.08) during Sectors 14–19 (2019 July 18 to 2019 December 24). After visually inspecting the target pixel files, we found that the host star had fallen outside of the CCD’s science image area in Sector 14. We thus performed our analysis with data from sectors 15–19.

The photometric observations for HD 191939 (see Figure 2) were processed through the Science Processing Operations Center (SPOC) pipeline, developed and maintained by the NASA Ames Research Center (Jenkins et al. 2016; Jenkins 2017).¹ The pipeline detected two tentative planetary signals in the combined transit search for Sectors 15 and 16. With the addition of Sectors 17–19, the MIT Quick Look Pipeline (QLP; Huang et al., in prep.) identified three recurring transit signals.

¹The SPOC pipeline searches for planetary transits by fitting an averaged Mandel & Agol (Mandel & Agol 2002) model to the light curve with non-linear limb-darkening coefficients as parametrized by Claret & Bloemen 2011 (Seader et al. 2013; Li et al. 2019).

The phase-folded light curves obtained with the SPOC transit parameters had a flat-bottomed shape, consistent with a planetary interpretation of the transits. Moreover, the two planet candidates passed all the SPOC and QLP standard validation diagnostics, including a search for secondary eclipses, differences in odd and even transits, and flux centroid offsets during transit (see Section 4.1).

We retrieved the SPOC-processed data from the *Mikulski Archive for Space Telescopes* (MAST).² In particular, we downloaded the Presearch Data Conditioning (PDC) light curves, and removed all the observations encoded as *NaN* or flagged as bad-quality points by the SPOC pipeline. From a total of 85282 photometric measurements (17848, 16812, 16945, 16612, and 17065 for Sectors 15–19, respectively), we identified a total of 4980 bad-quality data points, which we excluded from further analysis.

2.2. Ground-Based Photometry: Observatori Astronòmic Albanyà

As part of the TFOP follow-up program,³ we acquired 320 photometric exposures of HD 191939 on 2019 October 29 with the 0.4-m telescope at the Observatori Astronòmic Albanyà (OAA) in Catalonia (Spain). The host star was continuously observed for 398.8 minutes in the Cousins I_c filter using a CCD camera with a resolution of 3056×3056 pixels and a pixel scale of $0.72''$ per pixel. The science exposures were reduced with the AstroImageJ (AIJ) software (Collins et al. 2016).

2.3. Archival Spectroscopic Observations: SOPHIE

SOPHIE (Bouchy et al. 2009a; Perruchot et al. 2008) is a fiber-fed échelle spectrograph mounted on the 1.93-m telescope at the Observatoire de Haute Provence (OHP), in France. This instrument observed HD 191939 between 2007 September 27 and 2007 November 30 with a RV precision of $4\text{--}5 \text{ m s}^{-1}$ (e.g. Bouchy et al. 2009b, 2011). A total of five spectra were acquired with *SOPHIE*'s high-resolution mode, which provides a resolving power of $\lambda/\Delta\lambda \equiv R = 75,000$. The spectra have a median exposure time of 617 seconds and a median SNR per pixel at 550 nm of 59 (see Table 2).

We downloaded all the available observations of HD 191939 from the *SOPHIE* public archive (Moutaka et al. 2004). These included spectra reduced by the *Data Reduction Software v0.50*⁴ (DRS), as well as

Table 1. Stellar Properties of HD 191939.

Property	Value	Source
<i>Other Target Names</i>		
HD ID	191939	4
TOI ID	1339	-
TIC ID	269701147	1
2MASS ID	J20080574+6651019	2
Gaia DR2 ID	2248126315275354496	3
<i>Astrometric Properties</i>		
R.A. (J2015.5; h:m:s)	20:08:06.150	3
Dec (J2015.5; d:m:s)	+66:51:01.08	3
Parallax (mas)	18.706 ± 0.071	3
Distance (pc)	53.608 ± 0.073	1
$\mu_{\text{R.A.}}$ (mas y^{-1})	150.256 ± 0.044	3
$\mu_{\text{Dec.}}$ (mas y^{-1})	-63.909 ± 0.047	3
<i>Photometric Properties</i>		
<i>TESS</i> (mag)	8.292 ± 0.006	1
B (mag)	9.720 ± 0.038	5
V (mag)	8.97 ± 0.03	6
Gaia (mag)	8.7748 ± 0.0002	3
J (mag)	7.597 ± 0.029	2
H (mag)	7.215 ± 0.023	2
K_s (mag)	7.180 ± 0.021	2

References—(1) *TESS* Input Catalog Version 8 (TICv8) (Stassun et al. 2018). (2) Two Micron All Sky Survey (*2MASS*; Cutri et al. 2003). (3) *Gaia* DR2 (Brown et al. 2018). (4) Henry Draper Catalog (Cannon & Pickering 1993). (5) Tycho2 Catalog (Høg et al. 2000). (6) Hipparcos Catalog (van Leeuwen 2007).

NOTE— †: *Gaia* DR2 parallax and uncertainty corrected for a systematic offset of $+0.082 \pm 0.033$ mas, as described in Stassun & Torres 2018.

the cross-correlation functions (CCFs) determined by the DRS using a numerical mask for the G2 spectral type (Baranne et al. 1996). The CCFs were calculated over a ± 30 km/s velocity interval. The RV, full-width half maximum (FWHM), and contrast of each CCF were computed by the DRS by fitting a Gaussian function to the CCF profile. We extracted barycentric-corrected radial velocities, FWHM, and the bisector spans from the FITS headers of the CCFs (Coroller & Bouchy 2017). We found no correlations between either the RV measurements and the bisectors ($r = 0.51$, while the critical value of the Pearson correlation coefficient at the confidence level of 0.01 is $r_{3,0.01} = 0.96$), or between the RVs and the FWHM of the CCFs ($r = 0.43$). Such correla-

²<https://mast.stsci.edu/portal/Mashup/Clients/Mast/Portal.html>

³<https://tess.mit.edu/followup/>

⁴http://www.obs-hp.fr/guide/sophie/data_products.shtml.

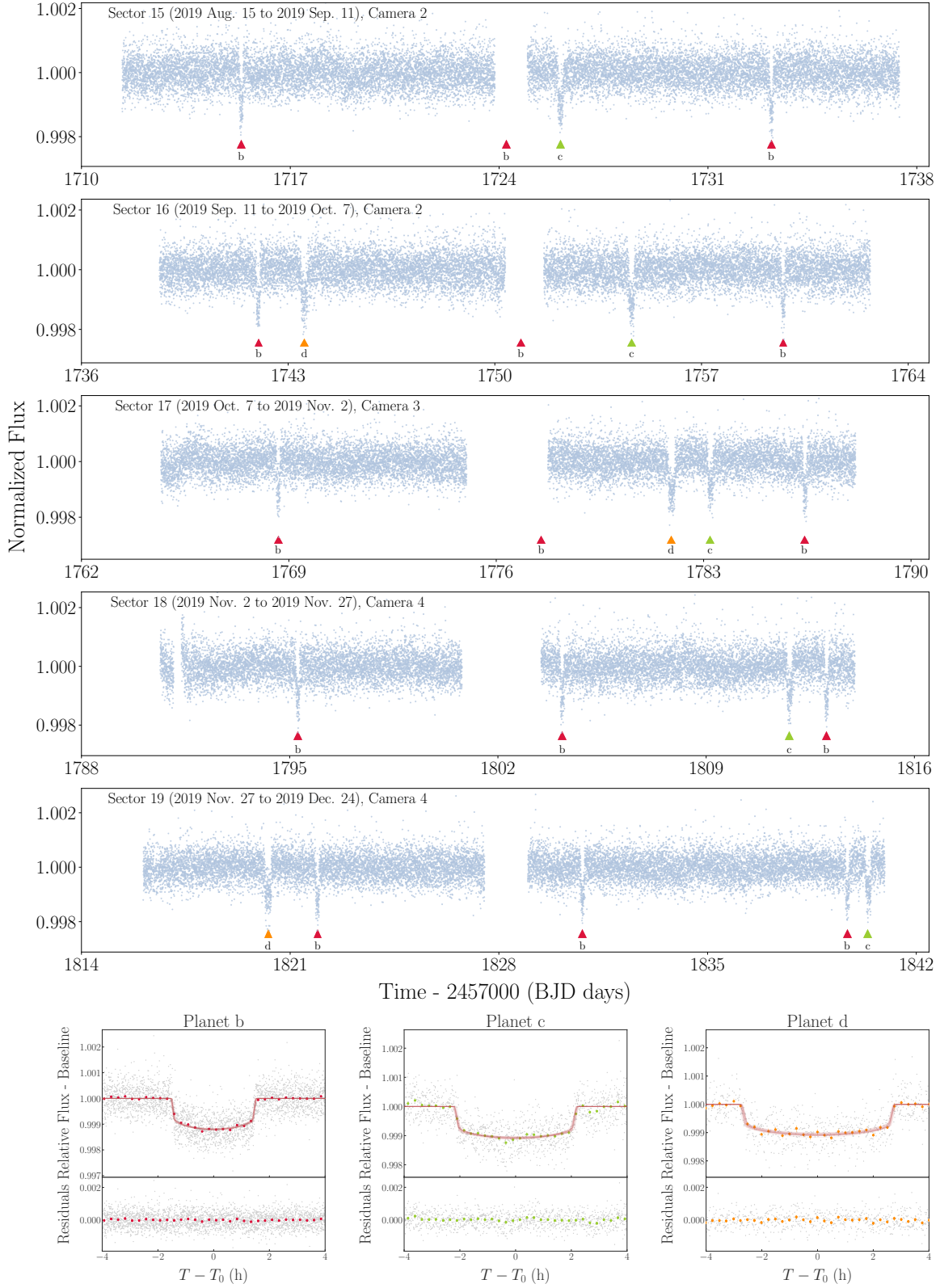


Figure 2. *Top:* The full *TESS* discovery light curve based on the 2-min. exposures from Sectors 15–19. The transits of planet b, c and d are shown in red, green, and orange, respectively. *Bottom:* *TESS* phase-folded light curves over the *allesfitter* best-fit periods and initial epochs (see Section 5). The grey points are the *TESS* 2-min. exposures, the colored circles are the data points binned over 15 minutes, and the red lines represent 20 posterior models drawn from the outcome of the final fit. The light curve residuals are shown in the bottom panel.

tions would have indicated astrophysical false positives, such as stellar spots or blends.

Table 2. *SOPHIE* RV measurements.

BJD _{UTC}	RV	Error
−2,450,000	(km s ^{−1})	(km s ^{−1})
4371.345	−9.237	0.002
4372.288	−9.249	0.001
4430.266	−9.213	0.001
4431.296	−9.218	0.001
4435.308	−9.232	0.001

2.4. Spectroscopic Follow-up

2.4.1. TRES Reconnaissance Spectroscopy

We obtained three spectra of HD 191939 between 2019 October 24 and 2019 November 5 using the fiber-fed Tillinghamast Reflector Echelle Spectrograph (TRES; Fürész 2008) on the 1.5-m telescope at the Fred Lawrence Whipple Observatory (Mt. Hopkins, Arizona). TRES covers the spectral range 3850–9100 Å and has a resolving power of $R = 44,000$. The TRES spectra have an average SNR per resolution element of 47 and were extracted as described in Buchhave et al. (2010).

2.4.2. LCO/NRES Reconnaissance Spectroscopy

We acquired three consecutive 20-minute optical exposures of HD 191939 on 2019 November 1 with the Network of Echelle Spectrographs (NRES; Siverd et al. 2016; Siverd et al. 2018), operated by Las Cumbres Observatory (LCO; Brown et al. 2013). The NRES is composed of four high-precision fiber-fed spectrographs with a resolving power of $R = 53,000$. The NRES spectra were stacked to remove cosmic rays and instrumental effects, resulting in a final SNR of 60.

2.5. Gemini/NIRI High Angular Resolution Imaging

We observed HD 191939 on 2019 November 8 with the Near InfraRed Imager (NIRI) at the Gemini North telescope (Hodapp et al. 2003). A total of nine Adaptive Optics (AO) images were collected in the Br γ filter, each with an exposure time of 2 seconds. The telescope was dithered in a grid pattern between each science exposure to construct the sky background frame and remove artifacts such as bad pixels and cosmic rays. Data were processed using a custom set of IDL codes with which we interpolated bad pixels, subtracted the sky background, flat corrected images, aligned the stellar position between frames, and co-added data. We determined the sensitivity to stellar companions by injecting artificial point spread functions into the data at a range of separation and angles, and scaling these until they could be

detected at 5σ . We are sensitive to stellar neighbors 5 magnitudes (8.4 magnitudes) fainter than HD 191939 at 200 mas (1"). Our sensitivity as a function of radius is shown in Figure 3 with a thumbnail image of HD 191939.

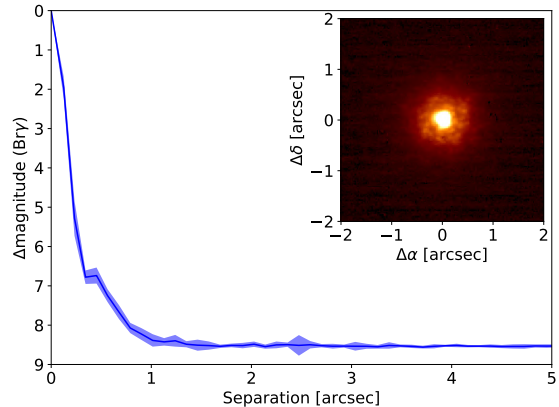


Figure 3. Sensitivity curve of our Gemini/NIRI AO images (solid black line). We are sensitive to companions with a contrast of 5 mag just 200 mas from the star. No visual sources are seen anywhere in the field-of-view. A thumbnail image of the target is inset.

3. HOST STAR CHARACTERIZATION

3.1. TRES Spectroscopy

We used the TRES spectra to measure the host star’s effective temperature T_{eff} , surface gravity $\log g$, metallicity $[\text{Fe}/\text{H}]$, and rotational velocity $v \sin i$ with the Stellar Parameter Classification (SPC; Buchhave et al. 2012) tool. The SPC software cross-correlates an observed spectrum against a grid of synthetic spectra based on the Kurucz atmospheric models (Kurucz 1992). The weighted average results are $T_{\text{eff}} = 5427 \pm 50$ K, $\log g = 4.44 \pm 0.10$ dex, $[\text{m}/\text{H}] = -0.16 \pm 0.08$ dex, and $v \sin i = 0.6 \pm 0.5$ km/s (see Table 3). From the TRES spectra, we also detected weak H-alpha absorption indicating that HD 191939 has a low stellar activity.

3.2. NRES Spectroscopy

We also constrained the stellar parameters from the NRES observations following the techniques presented in Petigura et al. (2017) and Fulton & Petigura (2018). In particular, we used the SpecMatch software (Petigura 2005; Petigura et al. 2017)⁵ to compare the observed spectrum of HD 191939 against synthetic spectra created by linearly interpolating the Coelho et al. (2005) grid of model spectra at arbitrary sets of T_{eff} , $\log g$, $[\text{Fe}/\text{H}]$, and $v \sin i$. We maximized the χ^2 -based likelihood via a Levenberg-Marquardt algorithm (William

⁵<https://github.com/petigura/specmatch-syn>

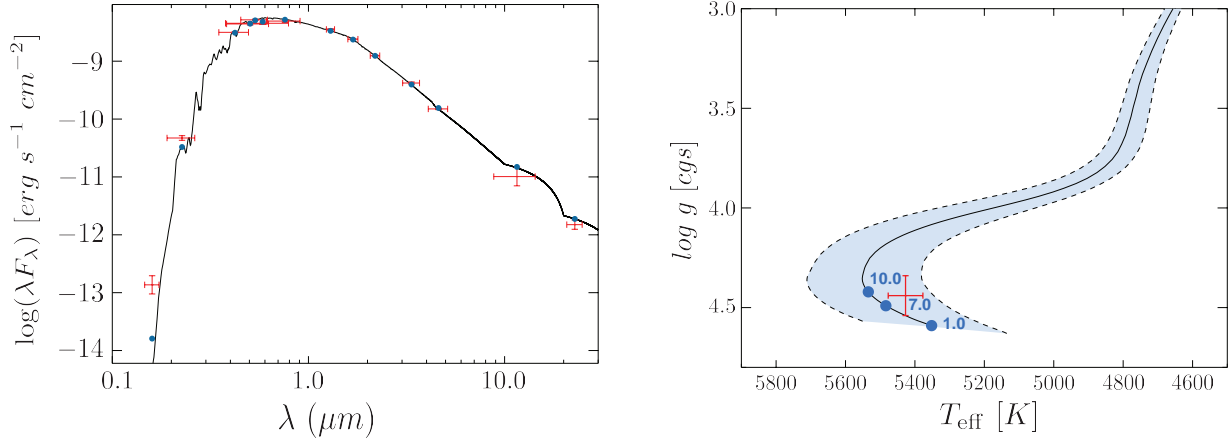


Figure 4. *Left:* Spectral energy distribution (SED). Red symbols are the observed photometric data, with the horizontal bars reflecting the effective width of the passband. Blue symbols are the model fluxes from the best-fit Kurucz atmosphere model (black). *Right:* H-R diagram. The black curve with blue swathe represents a Yonsei-Yale evolutionary model for the stellar mass and metallicity with their uncertainties. The blue dots label ages along the evolutionary track in Gyr. The red symbol represents the spectroscopically derived effective temperature and surface gravity with their uncertainties.

H. Press 1992) and found $T_{\text{eff}} = 5335 \pm 100 \text{ K}$, $\log g = 4.2 \pm 0.1 \text{ cgs}$, $[\text{Fe}/\text{H}] = -0.13 \pm 0.06 \text{ dex}$, and $v \sin i < 2 \text{ km/s}$.

3.3. Spectral Energy Distribution Analysis

We used the host star’s broadband Spectral Energy Distribution (SED) and its *Gaia* DR2 parallax to determine an empirical measurement of the stellar radius following procedures described in the literature (Stassun & Torres 2016; Stassun et al. 2017; Stassun et al. 2018). For this analysis, we retrieved the FUV and NUV fluxes from *GALEX*, the B_T and V_T magnitudes from *Tycho-2*, the J , H , and K_s magnitudes from *2MASS*, the W1–W4 magnitudes from *WISE*, and the G , G_{BP} , and G_{RP} magnitudes from *Gaia*. When taken in combination, the available photometry spans the full stellar SED over the wavelength range 0.15–22 μm (see Figure 4).

We performed a fit to the host star’s SED with the Kurucz stellar atmospheric models, placing priors on T_{eff} , $\log g$, and $[\text{Fe}/\text{H}]$ based on the SPC analysis of the TRES spectra. The remaining free parameter was the extinction (A_V), which we limited to the maximum line-of-sight value from the dust maps of Schlegel et al. (1998). The model fits the data well, with a reduced χ^2 of 1.9 and best-fit extinction of $A_V = 0.03 \pm 0.03$ (see Figure 4). We also integrated the model SED to obtain a bolometric flux at Earth of $F_{\text{bol}} = (7.81 \pm 0.18) \times 10^{-9} \text{ erg s}^{-1} \text{ cm}^{-2}$. Using F_{bol} , T_{eff} , and the *Gaia* DR2 parallax adjusted by $+0.08 \text{ mas}$ to account for the systematic offset reported by Stassun & Torres (2018), we determined a stellar radius of $R_* = 0.945 \pm 0.021 R_{\odot}$. We also estimated the stellar mass empirically. The eclipsing binary-based relations of Torres et al. (2010) yield $M_* = 0.92 \pm 0.06 M_{\odot}$, whereas the stellar surface gravity

and SED-based radius result in $M_* = 0.90 \pm 0.21 M_{\odot}$. Figure 4 shows the former in a Hertzsprung-Russell (H-R) diagram with an evolutionary track from the Yonsei-Yale models (Yi et al. 2001; Spada et al. 2013). These plots imply that the age of HD 191939 is $7 \pm 3 \text{ Gyr}$.

Finally, we used the spectroscopic $v \sin i$ and the SED-based radius to calculate a stellar rotation period of $P_{\text{rot}}/\sin i = 79 \pm 66 \text{ d}$, where the large uncertainty is driven by the large error on the spectroscopic $v \sin i$. There is no chromospheric activity measurement published, but this is consistent with the fact that the *Galex* photometry in the SED does not appear to be strongly in excess. This implies that the host star rotates no faster than the Sun. Moreover, given that HD 191939 is roughly solar mass, this also indicates that it is somewhat older than the Sun, as evidenced by the H-R diagram.

3.4. Independent Validation of Stellar Parameters

As an independent validation on the SPC/SED stellar parameters, we used the spectroscopic properties of HD 191939 derived from the TRES and NRES spectra to perform isochrone fitting with two stellar evolutionary models: the MESA Isochrones and Stellar Tracks database (MIST; Choi et al. 2016; Dotter 2016) as implemented by the *isochrones* (Morton 2015) and *isoclassify* (Huber et al. 2017) packages, and the Padova models (da Silva et al. 2006), accessible via the PARSEC v1.3 (Bressan et al. 2012) online tool.⁶

We ran *isochrones* and *isoclassify* with priors on the star’s photometric magnitudes, the corrected *Gaia*

⁶http://stev.oapd.inaf.it/cgi-bin/param_1.3

Table 3. Derived stellar properties for HD 191939.

Property	Value	Source	Reference spectra
R_* (R_\odot)	0.945 ± 0.021	SED	TRES
M_* (M_\odot)	0.92 ± 0.06	SED via Torres et al. 2010	TRES
Age (Gyr)	7 ± 3	SED	TRES
A_v	$0.03^{+0.03}_{-0.03}$	SED	TRES
F_{bol} (erg/s/cm ²)	$(7.81 \pm 0.18) \times 10^{-9}$	SED	TRES
T_{eff} (K)	5427 ± 50	SPC	TRES
$\log g$ (cgs)	4.40 ± 0.10	SPC	TRES
$[m/H]$ (dex)	-0.16 ± 0.08	SPC	TRES
$(v \sin i)_A$ (km s ⁻¹)	0.6 ± 0.5	SPC	TRES
L_* (L_\odot)	$0.69^{+0.01}_{-0.01}$	isochrones (MESA Isochrones & Stellar Tracks)	TRES
Distance (pc)	$53.48^{+0.19}_{-0.20}$	isochrones (MESA Isochrones & Stellar Tracks)	TRES
ρ (g cm ⁻³)	$1.64^{+0.88}_{-0.40}$	allesfitter	
Spectral type	G8V	Pecaut et al. (2012), Pecaut & Mamajek (2013)	
$W_{H\alpha}$ (Å)	1.259 ± 0.007	This Work	SOPHIE

DR2 parallax, and the best-fit spectroscopic parameters (T_{eff} , $\log g$, [Fe/H]) from either the TRES or NRES spectra. For `isoclassify`, we also accounted for extinction by incorporating the 3D dust map of Green et al. (2018), which covers most of the sky with a declination larger than -30° . The derived stellar parameters are consistent with the values presented in Table 3, regardless of the choice of reference spectra for the host star’s spectroscopic parameters. To assess whether this consistency was primarily due to the use of the MIST database, we also determined the stellar parameters with the Padova models with the `PARAM` v1.3 tool. For both the TRES and NRES spectroscopic parameters, the resulting stellar parameters agree with the `isochrones` and `isoclassify` predictions to within 1σ . In the rest of this paper, we adopt the results from Table 3 for our analysis of the HD 191939 system.

3.5. Chromospheric Activity Indicators

The *SOPHIE* spectra detailed in Section 2.3 indicate that HD 191939 is chromospherically inactive. As with the TRES spectra, the five *SOPHIE* spectra show the $H\alpha$ line in absorption. We measured the equivalent width of the $H\alpha$ line ($W_{H\alpha}$) in each spectrum using a 10-Å subsample centered on the vacuum wavelength of $H\alpha$ (6562.81 Å). We fitted a Voigt profile to the line and a linear trend to the continuum via least squares using `astropy` (Price-Whelan et al. 2018). We estimated the uncertainties by bootstrapping the model fit 100 times excluding a random 10% of the data points. The equivalent widths are consistent between the five spectra, and we measure a weighted-mean $H\alpha$ equivalent width of $W_{H\alpha} = 1.259 \pm 0.007$ Å. Visual inspection reveals no evidence of emission in the cores of the Ca II H and K

or $H\alpha$ lines. We conclude these factors indicate a lack of measurable chromospheric activity for HD 191939.

4. RULING OUT FALSE POSITIVES

The *Kepler* mission revealed that multiple-period transit-like events are more likely to be caused by true planets than by false positives (e.g. Latham et al. 2011; Lissauer et al. 2012). Despite such evidence, it is important to carefully inspect the observed transit signals to rule out false positive scenarios, including instrument systematics and contamination from nearby stars. In this section, we aim to validate the HD 191939 planet candidates with the SPOC Validation Tests (Section 4.1), the *TESS* photometry and archival/follow-up observations of HD 191939 (Section 4.2), and the statistical validation software `vespa` (Section 4.3).

4.1. SPOC Validation Diagnostics

The two planet candidates identified by the SPOC Data Validation Pipeline (referred to as planet “b” and “c” in our analysis; see Section 5) pass all the SPOC Data Validation (DV) tests (Twicken et al. 2018). We list these tests in the following paragraphs.

- An *Eclipsing Binary Discrimination Test* to search for weak secondary eclipses and compare the depth of odd and even transits. Planet b and c pass this diagnostic at 2σ , with no shallow secondaries around phase 0.5 and no odd/even transit depth variations.
- A *Ghost Diagnostic Test* primarily designed to rule out scattered light, nearby EBs, and instrumental noise as the source of the transit-like events. Both planet b and c pass this diagnostic within 2σ .
- A *Difference Image Centroid Offset Test* to determine if the location of the transit source is statisti-

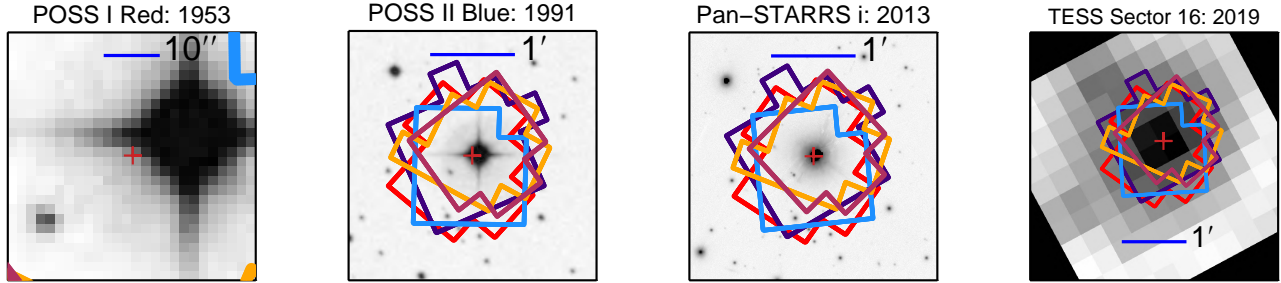


Figure 5. Archival images and *TESS* view of the field around HD 191939 from 1953 to 2019. The North points up and East is to the left. The red cross is the star’s current location, and the red, purple, light blue, orange, and maroon regions are the *TESS* photometric apertures for Sectors 15, 16, 17, 18, and 19 respectively. Due to the star’s proper motion, there is an offset between its present-day position and its location in the original POSS images.

cally consistent with the position of the target star. The offset distance for planet b in the combined Sector 15–16 SPOC analysis was less than $1.7''$ (0.33σ). For planet c, the maximum offset distance was less than $5.2''$ (0.94σ).

- A *Bootstrap Test* to assess the confidence level of the transit detection. Planet b and c pass this test with formal false alarm probabilities of 1.05×10^{-135} and 1.72×10^{-63} , respectively.

These DV tests were not applied to the third planet around HD 191939 (or planet “d”), as it was not detected by the SPOC pipeline in the combined transit search of *TESS* Sectors 15 and 16. Nevertheless, we independently verified the planetary nature of all three planet candidates with the analyses described in the following sections.

4.2. Observational Constraints

4.2.1. Archival Optical Images

The *TESS* detectors have a larger pixel scale than the *Kepler* telescope (*TESS*: $\sim 21''$, *Kepler*: $4''$), so photometric contamination from nearby astrophysical sources is more likely. To investigate this false positive scenario, we compared a *TESS* exposure of HD 191939 from Sector 16 to archival optical images taken in 1953, 1991, and 2013 by the first Palomar Observatory Sky Survey (POSS-I; Minkowski & Abell 1963), the second POSS (POSS-II; Reid et al. 1991) and the Pan-STARRS survey (Kaiser et al. 2002, 2010), respectively (see Figure 5). Due to HD 191939’s high proper motion, its present-day location appears unobscured in the archival images. Based on the POSS-I field-of-view, when HD 191939 was $\sim 11''$ away from its current sky position due to proper motion, we estimate that any $V \lesssim 19$ stars would have been clearly visible where HD 191939 is located today (e.g. a $V = 18.34$ source, identified as *Gaia* DR2

2248126310978337408, can be observed in the bottom left corner of the POSS-I view).

We performed a query of *Gaia* DR2 and *2MASS* catalogs centered on HD 191939, and used the SPOC reports to identify potential background sources around the host star. Within the central *TESS* pixel, there is only one *2MASS* source at $13.6''$ separation with a *TESS* magnitude of $T = 14.70$ mag (*2MASS* J20080397+6651023; TIC 269701145). However, this object is likely a spurious *2MASS* detection. First, *2MASS* artifacts are known to appear around bright stars, typically along their diffraction spikes (Stassun et al. 2018). Second, it was only observed in the J-band ($J = 14.2$ mag) and not in the H- and K-band (*2MASS* photometric quality flag of “AUU” and read flag of 0 for H- and K-bands; Cutri et al. 2003). Third, it is likely that *Gaia* DR2 would have detected this *2MASS* object provided it were real. With a *Gaia* magnitude difference of $\Delta G \sim 5.4$ mag relative to HD 191939,⁷ this source would have been within the observable parameter space of *Gaia* DR2’s contrast sensitivity curve (Brandeker & Cataldi 2019). Fourth, the Pan-STARRS images for HD 191939 in the *grizy* filters (e.g. see Figure 5) do not reveal any stellar objects near the position of the *2MASS* source. Given the high level of completeness of the Pan-STARRS survey down to its limiting magnitude (*grizy* ~ 22.4), it is thus improbable that the *2MASS* source is a true star. Finally, we can rule out the existence of this artifact with ground-based photometry (see Section 4.2.2). With all the aforementioned evidence, we conclude that the *2MASS* source is

⁷The *2MASS* detection lacks optical photometry, so there is no reliable way of estimating its *Gaia* magnitude. To calculate ΔG , we assume that this object is very red (e.g. a late M-dwarf) and take its *2MASS* J-band magnitude as an approximation of its *Gaia* magnitude. This conservative approach provides an estimate of how faint the *2MASS* source could be in the *Gaia* band-pass.

an instrumental artifact and could not have caused the transit-like events in the *TESS* light curve.

4.2.2. Ground-based Photometry

The OAA observations covered a full transit of the inner planet HD 191939 b, and showed a possible detection of a roughly 1200 ppm transit within a $13''$ photometric aperture. However, the data was not of sufficient quality to include in our global model fit. The longer periods of the planet candidates c and d have prevented successful ground-based photometric follow-up of their transits thus far.

To rule out nearby EBs, we conducted aperture photometry of all the visible sources within $2.5'$ of HD 191939 using a photometric aperture radius of $13''$. For each source, we employed the AIJ software to determine the root mean square error of its light curve, the predicted transit depth on the target star, and the resulting SNR. None of the sources considered in this analysis are bright enough to be a potential source of the *TESS* detection. Moreover, the OAA exposures show no evidence of the apparent *2MASS* instrumental artifact discussed in Section 4.2.1.

4.2.3. High-Resolution Imaging

It is important to check for stellar companions that can dilute the light curve, thus biasing the measured planetary radius or even be the source of false positives (Ciardi et al. 2015). To search for such companions, we examined the AO Gemini/NIRI final image (see Section 2.5) and found no visible stellar objects in the field-of-view around HD 191939 (Figure 3).

4.2.4. Archival Radial Velocities

We performed a joint fit to the *SOPHIE* RV and *TESS* observations with *allesfitter* to rule out possible substellar or stellar companions. Our combined fit finds a 3σ upper limit on the RV semi-amplitudes (K) of $K_b < 250$ m/s, $K_c < 300$ m/s, and $K_d < 250$ m/s for planet candidate b, c, and d, respectively. In contrast, a brown dwarf ($M \approx 13 M_{\text{Jup}}$) around HD 191939 would have $K_b = 620$ m/s and $K_c = 420$ m/s, and $K_d = 380$ m/s. The *SOPHIE* constraints lie well below these values, thus pointing to the planetary origin of the *TESS* transits.

In addition, we estimated the RV semi-amplitudes of the HD 191939 planet candidates from the standard RV equation (see Eq. 14 in Lovis & Fischer 2010) using the *allesfitter* orbital results (see Table 4) and the planets' masses estimated via the Mass-Radius (MR) relation of Wolfgang et al. 2016. With W16, we obtained $M_b = 13 \pm 2 M_{\oplus}$, $M_c = 12 \pm 2 M_{\oplus}$ and $M_d = 12 \pm 2 M_{\oplus}$ for planet candidates b, c, and

d, respectively. These masses correspond to RV semi-amplitudes ($K_b = 2.0 \pm 0.6$ m/s, $K_c = 1.0 \pm 0.4$ m/s, and $K_d = 1.0 \pm 0.4$ m/s) well within the range of the *SOPHIE* predictions. We also estimated the planetary masses and RV semi-amplitudes with the probabilistic MR relation of Chen & Kipping 2017 and found consistent results.

4.2.5. Ingress/Egress Test

We investigated whether a chance-aligned background or foreground EB could have caused the observed transits in the *TESS* light curve by placing an upper limit on the magnitude of a fully blended star. In a scenario of photometric contamination by blended light, the observed *TESS* transit depth (δ_{obs}) is given by:

$$\delta_{\text{obs}} \simeq \left(\frac{R_{p,\text{true}}}{R_*} \right)^2 \frac{F_{\text{blend}}}{F_{\text{blend}} + F_{\text{star}}} = \delta_{\text{true}} \frac{f}{1+f}, \quad (1)$$

where f is the flux ratio $f \equiv F_{\text{blend}}/F_{\text{star}}$, δ_{true} is the square of the true planet-to-star radius ratio in the absence of a blend, F_{blend} is the flux of the contaminating source, and F_{star} is the flux of HD 191939. Under the assumption of a central transit (i.e. $b = 0$), the observed transit depth δ_{obs} must satisfy (Eq. 21 in Seager & Mallen-Ornelas 2003):

$$\delta_{\text{obs}} \leq \delta_{\text{blend}} = \frac{\left(1 - \frac{t_F}{t_T}\right)^2}{\left(1 + \frac{t_F}{t_T}\right)^2}. \quad (2)$$

where t_F/t_T is the ratio of the full transit duration to the total transit duration t_T , which parametrizes the transit shape.⁸ For each planet candidate, we generated posterior probability distributions for the transit observables (i.e. δ_{obs} , t_T and t_F) by fitting the *TESS* light curve with the *allesfitter* package (Günther & Dyalan 2019, and in prep.),⁹ with the transits of the other planet candidates masked out. From these posteriors, we used Eq. 2 to estimate the maximum transit depth caused by the contaminating star (δ_{blend}) and determine a 3σ lower limit for t_F/t_T . For all planet candidates, we find $t_F/t_T \sim 0.90$, which suggests that the transits are box-shaped and thus less likely to be caused by a blend (Seager & Mallen-Ornelas 2003).

To determine the *TESS* magnitude of the faintest blended star (m_{blend}) capable of producing the observed transits in the *TESS* light curve, we calculated the

⁸The full transit duration is the time between ingress and egress (i.e. second to third contact). The total transit duration is the time between first and fourth contact.

⁹<https://github.com/MNGuenther/allesfitter>

flux ratio f with Eq. 1 and exploited the relation between stellar magnitudes and fluxes ($m_{\text{blend}} - m_{\text{star}} = -2.5 \log_{10} f$, where m_{star} is the *TESS* magnitude of HD 191939). Our analysis rules out blended stars fainter than 10.94, 11.73, and 9.57 at a 3σ level for planet candidate b, c, and d, respectively. Therefore, objects such as the spurious *2MASS* detection mentioned in Section 4.2.1 would automatically be discarded as the cause of the observed *TESS* transits. Other nearby stars within the *TESS* photometric aperture, such as TIC 269701151 (at $42.81''$, with $T = 15.63$) and TIC 269701155 (at $45.59''$, with $T = 15.84$), would not be bright enough either to produce the observed transit-like events. Consequently, the results of the ingress-egress test support the planetary nature of the *TESS* transit signals.

4.3. Statistical Validation of the HD 191939 System

The public software *vespa* (Morton 2012) uses Bayesian inference to calculate the probability that the *TESS* transits are compatible with astrophysical false positive scenarios. For each HD 191939 planet candidate, we ran *vespa* with the planet’s *TESS* phase-folded light curve, the *allesfitter* best-fit results for its orbital period, transit depth, and planet-to-star radius ratio (see Section 5, Table 4), and the host star’s *Gaia* DR2 coordinates, photometric magnitudes, and effective temperature, metallicity and surface gravity (see Table 1). We also included three observational constraints in our False Positive Probability (FPP) calculation, namely: the Gemini/NIRI contrast curve (see Section 2.5), a maximum blend radius of $1''$ based on the high-contrast sensitivity analysis from Section 4.2.3, and a maximum depth of a potential secondary eclipse of 5×10^{-5} . To calculate the latter, we masked out the observed transits of planet candidates b, c, and d on the full *TESS* discovery light curve, and estimated an upper limit on the shallowest transit depth that could be detected by running a Box-Fitting Least Squares algorithm with the public software *lightkurve* (Lightkurve Collaboration et al. 2018).

The resulting FPPs are less than 10^{-6} for all three planet candidates. Given that multiple transit-like signatures are more likely to be caused by genuine planets than by false positives, these FPPs must be enhanced by a “multiplicity boost” corresponding to ~ 15 for *TESS* targets (Guerrero 2020). For *TESS* planets with sizes up to $R_p = 6 R_{\oplus}$, this factor increases to 60. When applying the latter to our *vespa* results, we obtain FPPs lower than 10^{-7} for all planet candidates. We thus conclude that HD 191939 has three statistically validated *bona fide* planets and refer to them as planet b, c, and d in our subsequent analysis.

5. GLOBAL MODEL FIT: ORBITAL AND PLANETARY PARAMETERS

We first employed the publicly available Transit Least Squares (TLS) package (Hippke & Heller 2019) to look for planetary transits in the *TESS* light curve (see Figure 2).¹⁰ The TLS routine identified three possible planetary signals with periods of $P_b \approx 8.88$ days, $P_c \approx 28.58$ days, and $P_d \approx 38.35$ days, confirming the findings of the SPOC pipeline. Taking the TLS orbital periods and transit times as our initial guesses, we performed a preliminary fit to the *TESS* light curve with *allesfitter*. For our final fit to the *TESS* data, we assumed circular orbits and fitted a transit model with nine free parameters:

- the planet-to-star radius ratio, R_p/R_* , with uniform priors from 0 to 1,
- the sum of the planetary and stellar radii over the semi-major axis, $(R_p + R_*)/a_p$, with uniform priors from 0 to 1,
- the cosine of the orbital inclination, $\cos i_p$, with uniform priors from 0 to 1,
- the planetary orbital period, P_p with uniform priors from -0.05 to $+0.05$ days around the initial TLS guess,
- the initial transit time $T_{0;p}$, with uniform priors from -0.05 to $+0.05$ days around the initial TLS guess,
- a quadratic stellar limb-darkening function, sampled uniformly with the triangular sampling technique of Kipping (2013),
- white noise (jitter) scaling terms for the *TESS* light curve, and
- two GP hyper-parameters for the Matern-3/2 kernel: the characteristic amplitude $\ln \sigma$, and the timescale $\ln \rho$.

We used a Markov Chain Monte Carlo (MCMC) algorithm, implemented by the Affine-invariant MCMC ensemble sampler *emcee*, to determine the posterior probability distributions of all the model parameters. We initialized the MCMC with 200 walkers, each taking 130000 steps, and performed a burn-in of 10000 steps for each chain before calculating the final posterior distributions. The resulting phase-folded light curves are shown in Figure 2 together with the best transit model. The associated fit parameters and their 1σ uncertainties are listed in Table 4. Figure 10 shows the posteriors for these model parameters. To ensure MCMC convergence,

¹⁰The TLS searches for transit-like events in photometric light curves by fitting a physical transit model with ingress, egress, and stellar limb-darkening. This method enhances the detection efficiency by $\sim 10\%$ relative to the standard Box Least Squares algorithm (Kovács et al. 2002), which fits a boxcar function to the transit signatures.

Table 4. Final Model Fit Results.

Parameter	HD 191939 b	HD 191939 c	HD 191939 d
Radius Ratio, R_p/R_*	$0.03289^{+0.00093}_{-0.00071}$	$0.0315^{+0.0010}_{-0.0010}$	$0.03092^{+0.00059}_{-0.00059}$
Sum of Radii over Semi-major axis, $(R_* + R_p)/a_p$	$0.0507^{+0.0079}_{-0.0052}$	$0.0253^{+0.0042}_{-0.0042}$	$0.02082^{+0.00093}_{-0.00081}$
Cosine of Orbital Inclination, $\cos i_p$	$0.023^{+0.013}_{-0.015}$	$0.0150^{+0.0060}_{-0.0079}$	$0.0088^{+0.0019}_{-0.0020}$
Mid-transit Time, $T_{0;p}$ (BJD days)	$2458715.35547^{+0.00060}_{-0.00060}$	$2458726.0531^{+0.0012}_{-0.0012}$	$2458743.5505^{+0.0015}_{-0.0015}$
Orbital Period, P_p (days)	$8.880411^{+0.00069}_{-0.00069}$	$28.58060^{+0.00048}_{-0.00045}$	$38.3561^{+0.0012}_{-0.0012}$
Transit Depth, $\delta_{dil;p;TESS}$ (ppt)	$1.082^{+0.062}_{-0.046}$	$0.993^{+0.066}_{-0.066}$	$0.956^{+0.037}_{-0.037}$
Planet Radius, R_p (R_\oplus)	$3.37^{+0.13}_{-0.13}$	$3.22^{+0.15}_{-0.14}$	$3.16^{+0.11}_{-0.11}$
Semi-major Axis, a_p (AU)	$0.089^{+0.010}_{-0.012}$	$0.178^{+0.033}_{-0.026}$	$0.216^{+0.011}_{-0.011}$
Orbital Inclination, i_p (deg)	$88.66^{+0.84}_{-0.77}$	$89.14^{+0.46}_{-0.34}$	$89.49^{+0.12}_{-0.11}$
Impact Parameter, $b_{tra;p}$	$0.48^{+0.17}_{-0.28}$	$0.61^{+0.12}_{-0.27}$	$0.438^{+0.074}_{-0.088}$
Total Transit Duration, $T_{tot;p}$ (h)	$3.052^{+0.047}_{-0.031}$	$4.452^{+0.074}_{-0.066}$	$5.525^{+0.048}_{-0.048}$
Full Transit Duration, $T_{full;p}$ (h)	$2.800^{+0.031}_{-0.052}$	$4.020^{+0.080}_{-0.11}$	$5.114^{+0.055}_{-0.055}$
Equilibrium Temperature, $T_{eq;p}$ (K)	778^{+58}_{-41}	550^{+45}_{-45}	499^{+12}_{-11}
System Parameters in the <i>TESS</i> band-pass			
Limb darkening Coefficient 1, $u_{1;TESS}$	$0.50^{+0.14}_{-0.18}$		
Limb darkening Coefficient 2, $u_{2;TESS}$	$-0.08^{+0.25}_{-0.18}$		
Flux error, $\log \sigma_{TESS}$ (log rel.flux.)	$-7.6797^{+0.0027}_{-0.0027}$		
GP Characteristic Amplitude, $gp : \log \sigma(TESS)$	$-9.490^{+0.045}_{-0.045}$		
GP Timescale, $gp : \log \rho(TESS)$	$-1.33^{+0.13}_{-0.13}$		

we required that the auto-correlation time for each parameter be larger than 30 (Foreman-Mackey et al. 2013).

For each planet candidate, we also derived additional system parameters, including the planetary radius, semi-major axis, orbital inclination, and equilibrium temperature; the stellar density; the ratio of the stellar radius to the planet’s semi-major axis; and the light curve observables imprinted by the planet, namely, the diluted transit depth, the total transit duration, the transit duration between ingress and egress, and the impact parameter (see Table 4).

6. DISCUSSION

6.1. Dynamical Analysis

6.1.1. Orbital Stability

An important test of orbital architectures derived from observation includes an analysis of the long-term dynamical stability. Such tests have been performed for numerous systems to investigate the validity of Keplerian solutions and the dynamical evolution of the systems (e.g. Fabrycky et al. 2014). For the HD 191939 system, we performed N-body integrations using the Mercury Integrator Package (Chambers 1999). Based on the stellar parameters shown in Table 3 and the orbital properties listed in Table 4, we constructed a dynamical simulation that spanned 10^7 simulation years. The starting eccentricity for the planets was fixed to that of circular orbits using the estimated planetary masses from Section 4.2.4. Since the innermost planet has a

very short orbital period, we chose a conservative time step of 0.1 days to ensure accuracy of the dynamical model.

The results of the simulation indicate that the HD 191939 system is stable based on the observed orbital parameters. In addition, there is little interaction between the planets and their eccentricities remain below 0.01 for the duration of the simulation, resulting in minimal impact on insolation flux received by the planets that would affect climate (Kane & Torres 2017). In particular, the innermost planet retains a circular orbit since it is the most massive and primarily influenced by the host star.

6.1.2. Potential for Additional Planets

In Section 5, we conducted a full TLS search for a 3σ detection of a fourth planet and found no evidence of it. To probe our sensitivity limit and investigate the detectability of planets which might still be hidden in the *TESS* data due to low SNR or data gaps, we also performed a series of injection recovery tests. In particular, we generated planet signals and injected them into the *TESS* light curves using *allesfitter*, with planet sizes ranging from 0.8 to $4 R_\oplus$ and periods ranging from 2 to 160 days. We then searched for these signals with TLS, and recorded a signal to be recovered if the detected period matched any multiple of half the injected period to better than 5%. We find that *TESS*’ recovery is nearly complete for sub-Neptunes and super-Earths on

orbits less than ~ 80 days (Figure 6). However, the regime of Earth-sized planets and of small exoplanets on longer orbits remain open for exploration. This means that more transiting planets amenable to atmospheric characterization might still await to be discovered in future *TESS* sectors, while RV and TTV monitoring might unveil additional non-transiting companions.

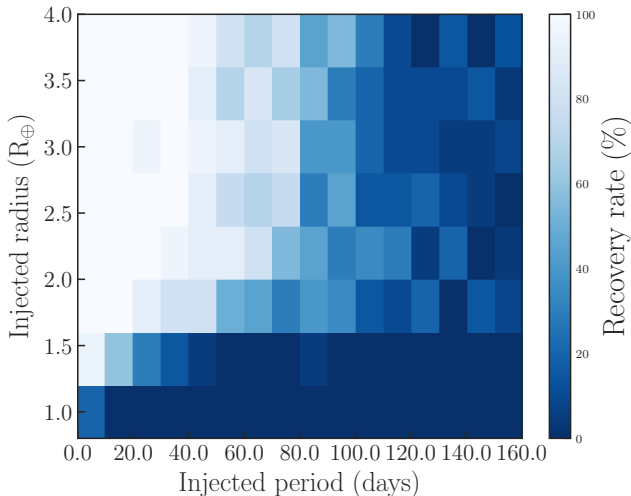


Figure 6. Injection-recovery test for simulated transits of small planets (0.8 to $4 R_{\oplus}$; y-axis) on periods between 2 to 160 days (x-axis). The color-coding shows the completeness of the recovery, with darker tones representing lower recovery rates. The *TESS* data for HD 191939 collected so far is near-complete for sub-Neptunes and super-Earths on orbits less than ~ 80 days, but the regime of the smallest and longest-period planets remains to be explored.

6.1.3. Transit Timing Variations

Transit timing variations (TTVs) in multi-planetary systems are caused by deviations from Keplerian orbits due to gravitational interactions between the planets (e.g. Agol et al. 2005; Holman & Murray 2005; Lithwick et al. 2012; Agol & Fabrycky 2018). These interactions can be used to verify the planetary nature of a transit signal and to characterize the planetary masses and the system’s orbital architecture. TTVs occur for systems in which pairs of planets orbit near mean motion resonances (MMR), where a ‘first-order MMR’ is defined as the period ratio being close to $P_{\text{in}}/P_{\text{out}} \approx (i - 1)/i$, whereby P_{in} and P_{out} are the periods of the inner and outer planets, and i is an integer. The planets’ mid-transit times then show sinusoidal variations with the ‘TTV super-period’, $P_{\text{TTV}} = |i/P_{\text{out}} - (i - 1)/P_{\text{in}}|^{-1}$.

For HD 191939, planets c and d are close to a first order MMR with a period ratio near 3:4. We would thus expect a TTV super-period of $P_{\text{TTV}} \sim 1500$ days – a

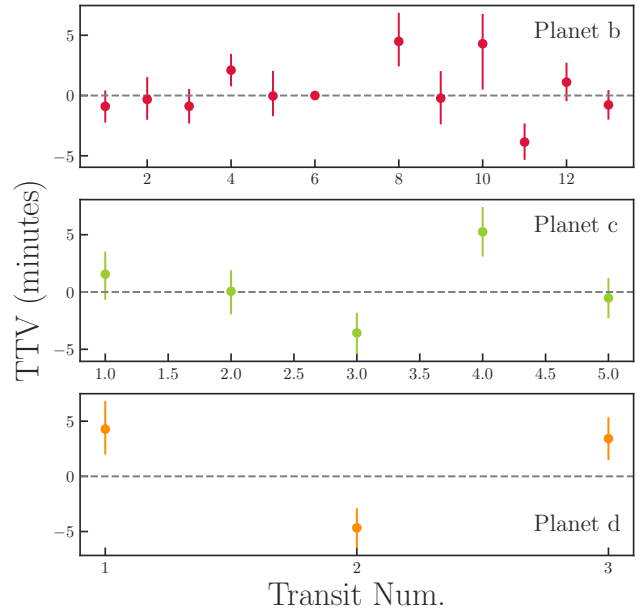


Figure 7. A search for TTV signals in the *TESS* data. From a free transit midtime fit, we find per-transit deviations from linear ephemerides, most notably for planets c and d. This could indicate a short-term TTV chopping signal.

factor of 10 longer than the span of our discovery data set (~ 150 days). This means we are currently only starting to sample the TTVs of this system, and are still in a regime where the linear period fits for planets c and d are likely biased. In contrast, the inner pair of planets, b and c, lie further off a second-order MMR with a period ratio near 1:3 and are thus expected to show much lower TTV amplitudes. In addition, short-timescale ‘chopping’ variations can occur when the planets are closest to another on their orbits (e.g. Deck & Agol 2015). These chopping TTVs typically occur on harmonics of the synodic timescale, $P_{\text{TTV;chopping}} = |1/P_{\text{out}} - 1/P_{\text{in}}|^{-1}$. For planets c and d, we expect this to happen on timescales of ~ 100 days, well within the available observation range.

We searched for evidence of TTVs by performing an *allesfitter* fit to the *TESS* light curve. For this analysis, we froze the initial epoch and orbital period, and fitted the rest of transit parameters described in Section 5, with the addition of a TTV parameter for each transit to allow for a shift in the mid-transit time. For independent confirmation, we also used the *exoplanet* (Foreman-Mackey et al. 2019) software and modeled the planetary orbits using the *TTVOrbit* class with Gaussian priors on the system parameters (from Section 5, Table 4). In each study, we placed uniform priors on the mid-transit time of each observed transit in the *TESS* light curve, centered on the expected mid-transit time

Table 5. Best sub-Neptune-sized ($R_p = 2 - 4R_\oplus$) planets for transmission spectroscopy. The predicted SNR is relative to that of GJ 143 b. We only show systems with measured masses and a relative error in host star radius, planet radius, and planet mass less than 30%. Data were retrieved from the NASA Exoplanet Archive on 2020 February 1.

Planet Name	Host Star	J magnitude (mag)	Relative SNR	Planet Radius (R_\oplus)	Discovery	Planet Radius Reference
1. GJ 436 b	GJ 436	6.9	8.34 ± 0.13	$3.96^{+0.05}_{-0.05}$	W. M. Keck	Knutson et al. (2011)
2. GJ 1214 b	GJ 1214	9.75	6.59 ± 0.23	$2.74^{+0.06}_{-0.05}$	MEarth	Kundurthy et al. (2011)
3. π Men c	HD 39091	4.86	4.78 ± 0.28	$2.06^{+0.03}_{-0.03}$	<i>TESS</i>	Gandolfi et al. (2018)
4. HD 97658 b	HD 97658	6.20	2.34 ± 0.12	$2.247^{+0.098}_{-0.095}$	W. M. Keck	Van Grootel et al. (2014)
5. HD 191939 b [†]	HD 191939	7.59	1.84 ± 0.25	$3.37^{+0.13}_{-0.13}$	<i>TESS</i>	This Work
6. HD 3167 c	HD 3167	7.54	1.63 ± 0.22	$2.85^{+0.24}_{-0.15}$	K2	Vanderburg et al. (2016)
7. HD 191939 c [†]	HD 191939	7.59	1.48 ± 0.26	$3.22^{+0.15}_{-0.14}$	<i>TESS</i>	This Work
8. GJ 9827 d	GJ 9827	7.98	1.45 ± 0.31	$2.022^{+0.046}_{-0.043}$	K2	Rice et al. (2019)
9. HD 191939 d [†]	HD 191939	7.59	1.42 ± 0.26	$3.16^{+0.11}_{-0.11}$	<i>TESS</i>	This Work
10. GJ 143 b	GJ 143	6.08	1.00 ± 0.18	$2.61^{+0.17}_{-0.16}$	<i>TESS</i>	Dragomir et al. (2019)

[†] Planetary masses estimated from the MR relation of Wolfgang et al. 2016.

from the global fit (Table 4) with a width of 1 hour. We determined convergence once the fits reached a chain length of at least 30 times the auto-correlation length for *allesfitter*, and a Gelman-Rubin statistic $\hat{R} < 1.001$ for *exoplanet* (Gelman & Rubin 1992).

While we find no significant evidence for the long-term super-period TTVs (as expected), we recognize a deviation of the transit midtimes from strictly linear ephemerides on shorter time scales (Figure 7). This could either be due to noise or hint towards a chopping signal. Most notably, the first transit of planet d arrives 4.3 ± 2.5 min. late, the second arrives 4.7 ± 1.8 min. early, and the third arrives 3.4 ± 1.9 min. late again. Similarly, the third transit of planet c arrives 3.6 ± 1.8 min. early and the fourth transit arrives 5.2 ± 2.1 min. late. Future *TESS* observations and ground-based photometric follow-up will be needed to search for the first conclusive evidence of a chopping signal and to constrain the presence of long-term super-period TTV trends. Initial analyses could be possible after the first full year of monitoring with *TESS*, when a quarter of the super-period will have been sampled.

6.2. Atmospheric Characterization Prospects

All three HD 191939 planets have the potential to be excellent transmission spectroscopy targets, contingent upon mass measurements. The equilibrium temperatures for a zero Bond albedo and efficient redistribution of heat to the nightside hemisphere are about 780 K, 550 K, and 500 K for planets b, c, and d, respectively. To estimate the amplitude of the atmospheres' transmission signal, we assumed hydrogen-dominated

compositions and used the predicted masses from the deterministic MR relationship of Wolfgang et al. 2016 ($M_b = 13 \pm 2 M_\oplus$, $M_c = M_d = 12 \pm 2 M_\oplus$). Under these conditions, a change in the planet radius corresponding to one pressure scale height (H) would result in a change in the transit depth of ~ 30 ppm for HD 191939 b and ~ 20 ppm for both HD 191939 c and HD 191939 d. At near-infrared wavelengths, absorption due to species such as H_2O and CH_4 can produce variations of a few pressure scale heights in the effective planetary radius, translating to transmission signals $\gtrsim 50 - 150$ ppm for all three planets. Given the brightness of the host star, this puts all three HD 191939 planets among the most favorable sub-Neptunes currently known for transmission spectroscopy.

To put HD 191939 in the broader context of the exoplanet population suitable for atmospheric characterization, we downloaded a list from the NASA Exoplanet Archive of all the confirmed planets with radii between $2-4 R_\oplus$. We also required these planets to have mass measurements and a relative error in planet mass, planet radius, and host star radius less than 30%. For planets with multiple measurements of a given parameter, we selected the reported value with the lowest total uncertainty. We then calculated the expected SNR of each planet as (Vanderburg et al. 2016):

$$SNR \propto \frac{R_p H \sqrt{F t_{14}}}{R_*^2}, \quad (3)$$

with the atmosphere's scale height given by $H = \frac{k_b T_{eq}}{\mu g}$. Above, F is the stellar flux, k_b is the Boltzmann's constant, T_{eq} is the planet's equilibrium temperature, μ is

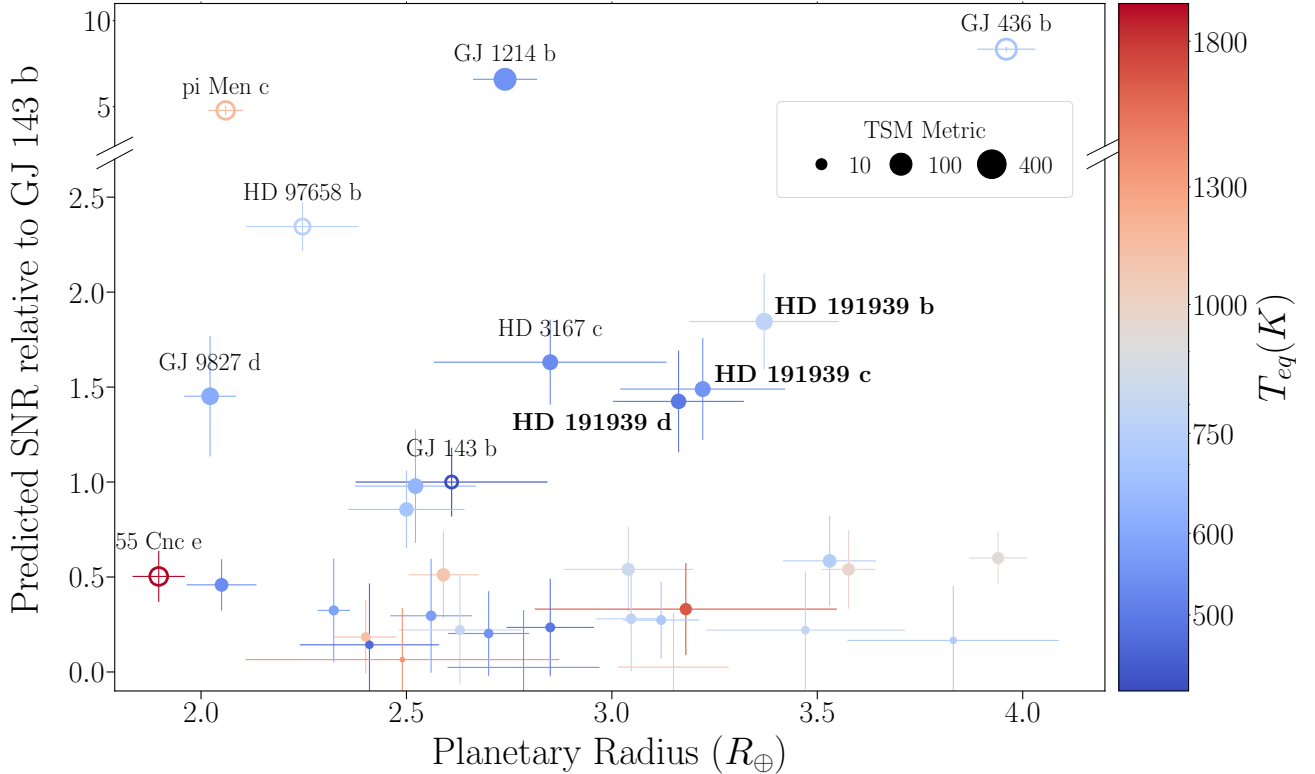


Figure 8. Predicted SNR relative to GJ 143 b for the population of confirmed sub-Neptunes ($R_p = 2 - 4 R_\oplus$), with the addition of 55 Cnc e ($R_p = 1.897^{+0.044}_{-0.046} R_\oplus$, Dai et al. 2019) for comparison. We only show systems with mass measurements and a relative error in host star radius, planet radius, and planet mass less than 30%. The color of the marker represents the planet’s equilibrium temperature, while its size indicates its TSM. Empty circles are used for planets with a host star brighter than the JWST brightness limit ($J \approx 7$ mag). Names are only displayed for the top 10 planets (see Table 5). Data were retrieved from the NASA Exoplanet Archive on 2020 February 1.

the atmospheric mean molecular weight, g is the planet’s surface gravity, and t_{14} is the transit duration (Seager 2010). We calculated the stellar flux from the host star’s H-band *2MASS* magnitude and set $\mu = 2.2$ atomic mass units (amu) under the assumption that sub-Neptunes have H_2 -dominated atmospheres. In the absence of t_{14} and T_{eq} values listed on the Exoplanet Archive, we computed these parameters with Eq. 16 and Eq. 2.27 in Seager & Mallen-Ornelas 2003 and Seager 2010, respectively. For these calculations, we assumed a zero Bond albedo ($A_B = 0$) and full heat redistribution over the planet’s surface ($f' = 1/4$).

Table 5 ranks the HD 191939 system relative to the known sub-Neptunes with measured masses. All three HD 191939 planets are among the best ten candidates for atmospheric characterization, with HD 191939 b offering the highest SNR predictions for transmission spectroscopy, followed by HD 191939 c and HD 191939 d. However, it is necessary to measure the planetary masses before this can be confirmed.

Figure 8 shows the HD 191939 planets in the context of the sub-Neptunes considered in this study. It also

illustrates the planets’ equilibrium temperatures and their Transmission Spectroscopy Metric (TSM; Kempton et al. 2018).¹¹ From an anticipated SNR perspective, HD 3167 c (Vanderburg et al. 2016) offers a useful point of comparison. The latter is a $2.9 R_\oplus$ planet orbiting a bright ($J = 7.5$ mag) KOV host star, with an equilibrium temperature of 600 K and a transmission signal amplitude of ~ 20 ppm for a $1H$ change in effective planet radius. Given the similar brightness of the HD 3167 and HD 191939 host stars, this means that HD 191939 c and HD 191939 d could be comparably suited for transmission spectroscopy, and that HD 191939 b could be more favorable than HD 3167 c, owing to its higher equilibrium temperature and thus larger atmospheric scale height.

¹¹The TSM predicts the expected transmission spectroscopy SNR of a 10-hour observing campaign with JWST/NIRISS under the assumption of a fixed MR relationship, cloud-free atmospheres, and the same atmospheric composition for all planets of a given type.

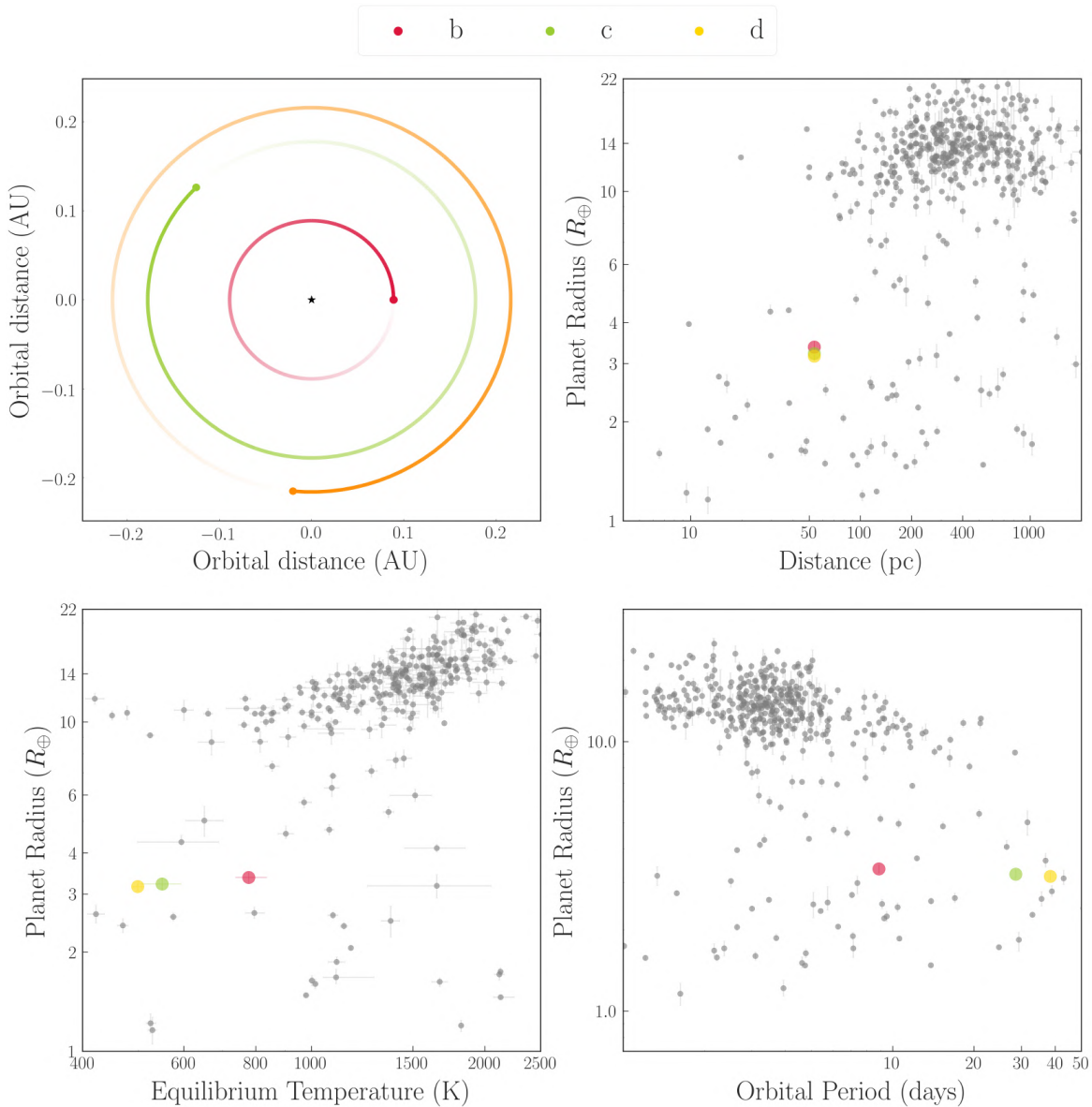


Figure 9. Planets b, c and d are shown in red, green, and orange, respectively. *Top left:* Top-down view of the HD 191939 system, with the sizes of the planets drawn to scale. *Top right, bottom left, and bottom right:* Planet radius as a function of distance, equilibrium temperature, and orbital period, respectively. The systems shown in these views have mass measurements for at least one of their planets and a relative error in host star radius, planet radius, and planet mass better than 30%. Data were retrieved from the NASA Exoplanet Archive on 2020 February 1.

6.3. HD 191939 in Context

The Sun-like star HD 191939 hosts three transiting sub-Neptunes in a compact orbital configuration. This system is a promising candidate for detailed characterization, as evidenced by Figure 1 and Figure 9.

First, the host star’s brightness and proximity make HD 191939 an excellent target for future photometric follow-up. In the context of single- and multi-planetary systems for which mass measurements are available and the relative error in host star radius, planet mass, and

planet radius is less than 30% (see Figure 1), HD 191939 ($K = 7.16$ mag) is one of the brightest and closest ($d \leq 100$ pc) stars with multiple transiting planets. When compared to multis with at least three planets, HD 191939 is the fifth brightest star in the K-band after HD 219134 (Motalebi et al. 2015), 55 Cnc (Butler et al. 1997), GJ 357 (Luque et al. 2019), and HD 3167 (Vanderburg et al. 2016), with GJ 9827 (Niraula et al.

2017) trailing behind.¹² Of all these multis, however, GJ 9827 is the only one with three *transiting* planets. As of February 2020, HD 219134 and HD 3167 are known to host two transiting planets, while 55 Cnc and GJ 357 only have one.

In addition, HD 191939 is one of the closest and brightest Sun-like (G-type) stars with transiting planets. When compared to the systems shown in Figure 1, HD 191939 is the second closest/brightest star known to date after the G8V dwarf 55 Cnc. However, the latter has a J-magnitude of $J = 4.76$ mag (Stassun et al. 2018), which exceeds the JWST brightness limit of $J \approx 7$ mag (Beichman et al. 2014). In contrast, HD 191939 has a J-magnitude of $J = 7.59$ mag, so it will likely be more suitable for future observations with JWST. With respect to the population of single- and multi-planetary systems specifically composed of G-type stellar hosts and transiting sub-Neptunes ($R_p = 2 - 4R_\oplus$), HD 191939 is the second closest/brightest system discovered to date after HD 39091 (Huang et al. 2018). With the addition of 55 Cnc, for which the lowest and highest reported radius of its transiting planet (55 Cnc e) are, respectively, $R_p = 1.897^{+0.044}_{-0.046} R_\oplus$ (Dai et al. 2019) and $R_p = 2.173^{+0.097}_{-0.098} R_\oplus$ (Endl et al. 2012), HD 191939 ranks third.

HD 191939 is also a valuable target for follow-up photometry due to its location in the northern ecliptic hemisphere sky. Indeed, HD 191939 lies in a region where 10 *TESS* sectors overlap (Sectors 14–19, 21–22, 24–25), hence enabling a long *TESS* observational baseline. In particular, HD 191939 will have been monitored for about 10 months once *TESS* Sector 25 (2020 June 08) concludes and will be re-observed for an additional ~ 10 months during the *TESS* extended mission. As a result, HD 191939 will have a photometric baseline of almost ~ 3 years. Such a long time span will facilitate a variety of dynamical studies, including a refinement of the system’s transit ephemerides and a search for additional planetary companions via TTV analyses. Moreover, the proximity of the host star (~ 54 pc) will also facilitate searches for massive planetary companions on wide orbits (~ 1 AU) using *Gaia* astrometry (Perryman et al. 2014).

From the perspective of ground-based RV follow-up, HD 191939 is also an excellent target for future observations due to its brightness, slow rotation, and lack of

measurable chromospheric activity. Future RV monitoring, with spectrographs such as CARMENES (Quirrenbach et al. 2014, 2018) or HARPS (Cosentino et al. 2012), will soon enable precise measurements of the planets’ masses. A RV monitoring campaign carried out by the CARMENES consortium is currently ongoing to confirm and further characterize the HD 191939 planets. Due to the system’s complex orbital architecture, however, a large number of measurements will be needed to accurately constrain the physical properties of the system.

With three temperate sub-Neptunes, HD 191939 is also a prime system for atmospheric studies with present (e.g. the *Hubble Space Telescope*) and future (e.g. JWST, Extreme Large Telescopes) facilities. With a high SNR for multi-wavelength transmission spectroscopy (see Table 5), the three HD 191939 planets are among the best ten candidates for atmospheric characterization studies (Figure 8). Their final suitability will be confirmed when mass measurements become available.

Furthermore, of all the planetary systems listed in Table 5, HD 191939 is the only multi for which the atmospheric characterization of *all* its transiting planets is possible. The other multi-planetary systems (i.e. HD 39091 and GJ 143) only have one planet amenable to transmission spectroscopy. With regard to JWST in particular, HD 191939 will be visible for more than 200 days per year due to its high ecliptic latitude (see Table 1).¹³ Future observations will provide an opportunity to study the atmospheres, interiors, and habitability conditions of the HD 191939 planets.

Finally, the multi-planetary nature of HD 191939 provides a fantastic opportunity to perform comparative exoplanetology. By studying the physical and orbital characteristics of HD 191939 and comparing them to the properties of the known population of multi-transiting planetary systems, we may gain insight into the distribution and occurrence rates of planets across a wide range of masses and radii, as well as into the formation and evolution of planetary architectures.

7. CONCLUSION

We have presented the *TESS* discovery of three sub-Neptune-sized planets around the nearby, bright Sun-like star HD 191939. We have confirmed the planetary nature of each planet candidate, both empirically through optical, photometric, and spectroscopic observations, and statistically via the public FPP implemen-

¹²The stars’ K-band magnitude, as well as the number of planets they host (n_{pl}), are: *HD 219134*: $K = 3.26$ mag, $n_{\text{pl}} = 6$. *55 Cnc*: $K = 4.01$ mag, $n_{\text{pl}} = 5$. *GJ 357*: $K = 6.47$ mag, $n_{\text{pl}} = 3$. *HD 3167*: $K = 7.06$ mag, $n_{\text{pl}} = 3$. *GJ 9827*: $K = 7.19$ mag, $n_{\text{pl}} = 3$.

¹³<https://jwst-docs.stsci.edu/jwst-observatory-hardware/jwst-target-viewing-constraints>

tation *vespa*. Upon refining the stellar parameters of HD 191939 reported by TICv8, we have derived the physical and orbital properties of the system with an *allesfitter* fit to the *TESS* discovery light curve. In addition, we have conducted a dynamical study of the HD 191939 planets, which indicates they are in a compact and stable orbital configuration consistent with circular orbits. Finally, we have demonstrated that the system is a promising target for precise photometric and RV follow-up as well as for future atmospheric characterization studies.

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Facilities: TESS, FLWO: 1.5m (TRES), LCO: 1m (NRES), OHP: 1.93m (SOPHIE), Gemini/NIRI, OAA: 0.4m.

Software: Python (G. van Rossum 1995), numpy (Oliphant 2006), scipy (Virtanen et al. 2020), matplotlib (Hunter 2007), astropy (Price-Whelan et al. 2018), pandas (McKinney 2010), allesfitter (Günther & Daylan 2019, and in prep.) emcee (Foreman-Mackey

et al. 2013), corner (Foreman-Mackey 2016), tqdm (doi:10.5281/zenodo.1468033), lightkurve (Lightkurve Collaboration et al. 2018), Transit Least Squares (Hipke & Heller 2019), vespa (Morton 2012), isochrones (Morton 2015), isoclassify (Huber et al. 2017), forecaster (Chen & Kipping 2017), exoplanet (Foreman-Mackey et al. 2019), starry (Luger et al. 2019), pymc3 (Salvatier et al. 2016), theano (Theano Development Team 2016), rebound (Rein & Liu 2012).

REFERENCES

- Agol, E., & Fabrycky, D. C. 2018, *Transit-Timing and Duration Variations for the Discovery and Characterization of Exoplanets*, 7, doi: [10.1007/978-3-319-55333-7_7](https://doi.org/10.1007/978-3-319-55333-7_7)
- Agol, E., Steffen, J., Sari, R., & Clarkson, W. 2005, *MNRAS*, 359, 567, doi: [10.1111/j.1365-2966.2005.08922.x](https://doi.org/10.1111/j.1365-2966.2005.08922.x)
- Baranne, A., Queloz, D., Mayor, M., et al. 1996, *A&AS*, 119, 373
- Beichman, C., Benneke, B., Knutson, H., et al. 2014, *Observations of Transiting Exoplanets with the James Webb Space Telescope (JWST)*, Publications of the Astronomical Society of the Pacific (PASP), December 2014. <https://arxiv.org/abs/1411.1754>
- Bouchy, F., Hébrard, G., Udry, S., et al. 2009a, *A&A*, 505, 853, doi: [10.1051/0004-6361/200912427](https://doi.org/10.1051/0004-6361/200912427)
- . 2009b, *A&A*, 505, 853, doi: [10.1051/0004-6361/200912427](https://doi.org/10.1051/0004-6361/200912427)
- Bouchy, F., Hébrard, G., Delfosse, X., et al. 2011, in *EPSC-DPS Joint Meeting 2011*, Vol. 2011, 240
- Brandeker, A., & Cataldi, G. 2019, *Astronomy & Astrophysics*, 621, A86, doi: [10.1051/0004-6361/201834321](https://doi.org/10.1051/0004-6361/201834321)
- Bressan, A., Marigo, P., Girardi, L., et al. 2012, *MNRAS*, 427, 127, doi: [10.1111/j.1365-2966.2012.21948.x](https://doi.org/10.1111/j.1365-2966.2012.21948.x)
- Brown, A. G. A., Vallenari, A., Prusti, T., et al. 2018, *Astronomy & Astrophysics*, 616, A1, doi: [10.1051/0004-6361/201833051](https://doi.org/10.1051/0004-6361/201833051)
- Brown, T. M., Baliber, N., Bianco, F. B., et al. 2013, *Publications of the Astronomical Society of the Pacific*, 125, 1031–1055, doi: [10.1086/673168](https://doi.org/10.1086/673168)
- Buchhave, L. A., Bakos, G. Á., Hartman, J. D., et al. 2010, *ApJ*, 720, 1118, doi: [10.1088/0004-637X/720/2/1118](https://doi.org/10.1088/0004-637X/720/2/1118)
- Buchhave, L. A., Latham, D. W., Johansen, A., et al. 2012, *Nature*, 486, 375, doi: [10.1038/nature11121](https://doi.org/10.1038/nature11121)
- Butler, R. P., Marcy, G. W., Williams, E., Hauser, H., & Shirts, P. 1997, *ApJL*, 474, L115, doi: [10.1086/310444](https://doi.org/10.1086/310444)
- Cannon, A. J., & Pickering, E. C. 1993, *VizieR Online Data Catalog*, III/135A
- Chambers, J. E. 1999, *MNRAS*, 304, 793, doi: [10.1046/j.1365-8711.1999.02379.x](https://doi.org/10.1046/j.1365-8711.1999.02379.x)
- Chen, J., & Kipping, D. 2017, *ApJ*, 834, 17, doi: [10.3847/1538-4357/834/1/17](https://doi.org/10.3847/1538-4357/834/1/17)
- Choi, J., Dotter, A., Conroy, C., et al. 2016, *ApJ*, 823, 102, doi: [10.3847/0004-637X/823/2/102](https://doi.org/10.3847/0004-637X/823/2/102)
- Ciardi, D. R., Beichman, C. A., Horch, E. P., & Howell, S. B. 2015, *ApJ*, 805, 16, doi: [10.1088/0004-637X/805/1/16](https://doi.org/10.1088/0004-637X/805/1/16)
- Claret, A., & Bloemen, S. 2011, *A&A*, 529, A75, doi: [10.1051/0004-6361/201116451](https://doi.org/10.1051/0004-6361/201116451)
- Cloutier, R., Doyon, R., Bouchy, F., & Hébrard, G. 2018, *The Astronomical Journal*, 156, 82, doi: [10.3847/1538-3881/aacea9](https://doi.org/10.3847/1538-3881/aacea9)
- Coelho, P., Barbuy, B., Meléndez, J., Schiavon, R. P., & Castilho, B. V. 2005, *A&A*, 443, 735, doi: [10.1051/0004-6361:20053511](https://doi.org/10.1051/0004-6361:20053511)
- Collins, K. A., Kielkopf, J. F., Stassun, K. G., & Hessman, F. V. 2016, *arXiv e-prints*, arXiv:1601.02622. <https://arxiv.org/abs/1601.02622>
- Coroller, H. L., & Bouchy, F. 2017, *SOPHIE: Documentation pour les utilisateurs*
- Cosentino, R., Lovis, C., Pepe, F., et al. 2012, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 8446, Harps-N: the new planet hunter at TNG, 84461V, doi: [10.1117/12.925738](https://doi.org/10.1117/12.925738)
- Cutri, R. M., Skrutskie, M. F., van Dyk, S., et al. 2003, *VizieR Online Data Catalog*, II/246
- da Silva, L., Girardi, L., Pasquini, L., et al. 2006, *A&A*, 458, 609, doi: [10.1051/0004-6361:20065105](https://doi.org/10.1051/0004-6361:20065105)
- Dai, F., Masuda, K., Winn, J. N., & Zeng, L. 2019, *ApJ*, 883, 79, doi: [10.3847/1538-4357/ab3a3b](https://doi.org/10.3847/1538-4357/ab3a3b)
- Deck, K. M., & Agol, E. 2015, *ApJ*, 802, 116, doi: [10.1088/0004-637X/802/2/116](https://doi.org/10.1088/0004-637X/802/2/116)
- Dotter, A. 2016, *The Astrophysical Journal Supplement Series*, 222, 8, doi: [10.3847/0067-0049/222/1/8](https://doi.org/10.3847/0067-0049/222/1/8)
- Dragomir, D., Teske, J., Günther, M. N., et al. 2019, *ApJL*, 875, L7, doi: [10.3847/2041-8213/ab12ed](https://doi.org/10.3847/2041-8213/ab12ed)

- Endl, M., Robertson, P., Cochran, W. D., et al. 2012, *ApJ*, 759, 19, doi: [10.1088/0004-637X/759/1/19](https://doi.org/10.1088/0004-637X/759/1/19)
- Fabrycky, D. C., Lissauer, J. J., Ragozzine, D., et al. 2014, *ApJ*, 790, 146, doi: [10.1088/0004-637X/790/2/146](https://doi.org/10.1088/0004-637X/790/2/146)
- Foreman-Mackey, D. 2016, *The Journal of Open Source Software*, 1, 24, doi: [10.21105/joss.00024](https://doi.org/10.21105/joss.00024)
- Foreman-Mackey, D., Czekala, I., Luger, R., et al. 2019, *dfm/exoplanet v0.2.3*, doi: [10.5281/zenodo.1998447](https://doi.org/10.5281/zenodo.1998447)
- Foreman-Mackey, D., Hogg, D. W., Lang, D., & Goodman, J. 2013, *PASP*, 125, 306, doi: [10.1086/670067](https://doi.org/10.1086/670067)
- Fressin, F., Torres, G., Charbonneau, D., et al. 2013, *The Astrophysical Journal*, 766, 81, doi: [10.1088/0004-637x/766/2/81](https://doi.org/10.1088/0004-637x/766/2/81)
- Fulton, B. J., & Petigura, E. A. 2018, *AJ*, 156, 264, doi: [10.3847/1538-3881/aae828](https://doi.org/10.3847/1538-3881/aae828)
- Fürész, G. 2008, PhD thesis, University of Szeged, Hungary
- G. van Rossum. 1995, Python tutorial, Technical Report CS-R9526, Tech. rep., Amsterdam
- Gandolfi, D., Barragán, O., Livingston, J. H., et al. 2018, *A&A*, 619, L10, doi: [10.1051/0004-6361/201834289](https://doi.org/10.1051/0004-6361/201834289)
- Gelman, A., & Rubin, D. B. 1992, *Statistical Science*, 7, 457, doi: [10.1214/ss/1177011136](https://doi.org/10.1214/ss/1177011136)
- Green, G. M., Schlafly, E. F., Finkbeiner, D., et al. 2018, *Monthly Notices of the Royal Astronomical Society*, 478, 651–666, doi: [10.1093/mnras/sty1008](https://doi.org/10.1093/mnras/sty1008)
- Guerrero, N. 2020, in *American Astronomical Society Meeting Abstracts*, Vol. 52, American Astronomical Society Meeting Abstracts, 327.03
- Günther, M. N., & Daylan, T. 2019, *allesfitter: Flexible star and exoplanet inference from photometry and radial velocity*. <http://ascl.net/1903.003>
- Günther, M. N., Pozuelos, F. J., Dittmann, J. A., et al. 2019, *Nature Astronomy*, 420, doi: [10.1038/s41550-019-0845-5](https://doi.org/10.1038/s41550-019-0845-5)
- Hippke, M., & Heller, R. 2019, *A&A*, 623, A39, doi: [10.1051/0004-6361/201834672](https://doi.org/10.1051/0004-6361/201834672)
- Hodapp, K. W., Jensen, J. B., Irwin, E. M., et al. 2003, *PASP*, 115, 1388, doi: [10.1086/379669](https://doi.org/10.1086/379669)
- Høg, E., Fabricius, C., Makarov, V. V., et al. 2000, *A&A*, 355, L27
- Holman, M. J., & Murray, N. W. 2005, *Science*, 307, 1288, doi: [10.1126/science.1107822](https://doi.org/10.1126/science.1107822)
- Howard, A. W., Marcy, G. W., Johnson, J. A., et al. 2010, *Science*, 330, 653–655, doi: [10.1126/science.1194854](https://doi.org/10.1126/science.1194854)
- Huang, C. X., Burt, J., Vanderburg, A., et al. 2018, *ApJL*, 868, L39, doi: [10.3847/2041-8213/aaef91](https://doi.org/10.3847/2041-8213/aaef91)
- Huber, D., Zinn, J., Bojsen-Hansen, M., et al. 2017, *ApJ*, 844, 102, doi: [10.3847/1538-4357/aa75ca](https://doi.org/10.3847/1538-4357/aa75ca)
- Hunter, J. D. 2007, *Computing in Science & Engineering*, 9, 90, doi: [10.1109/MCSE.2007.55](https://doi.org/10.1109/MCSE.2007.55)
- Jenkins, J. M. 2017, *Kepler Data Processing Handbook*, Tech. rep.
- Jenkins, J. M., Twicken, J. D., McCauliff, S., et al. 2016, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 9913, The TESS science processing operations center, 99133E, doi: [10.1117/12.2233418](https://doi.org/10.1117/12.2233418)
- Kaiser, N., Aussel, H., Burke, B. E., et al. 2002, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 4836, Pan-STARRS: A Large Synoptic Survey Telescope Array, ed. J. A. Tyson & S. Wolff, 154–164, doi: [10.1117/12.457365](https://doi.org/10.1117/12.457365)
- Kaiser, N., Burgett, W., Chambers, K., et al. 2010, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 7733, The Pan-STARRS wide-field optical/NIR imaging survey, 77330E, doi: [10.1117/12.859188](https://doi.org/10.1117/12.859188)
- Kane, S. R., & Torres, S. M. 2017, *AJ*, 154, 204, doi: [10.3847/1538-3881/aa8fce](https://doi.org/10.3847/1538-3881/aa8fce)
- Kempton, E. M.-R., Bean, J. L., Louie, D. R., et al. 2018, *Publications of the Astronomical Society of the Pacific*, 130, 114401, doi: [10.1088/1538-3873/aad6f6](https://doi.org/10.1088/1538-3873/aad6f6)
- Kipping, D. M. 2013, *Monthly Notices of the Royal Astronomical Society*, 435, 2152–2160, doi: [10.1093/mnras/stt1435](https://doi.org/10.1093/mnras/stt1435)
- Knutson, H. A., Madhusudhan, N., Cowan, N. B., et al. 2011, *ApJ*, 735, 27, doi: [10.1088/0004-637X/735/1/27](https://doi.org/10.1088/0004-637X/735/1/27)
- Kovács, G., Zucker, S., & Mazeh, T. 2002, *A&A*, 391, 369, doi: [10.1051/0004-6361:20020802](https://doi.org/10.1051/0004-6361:20020802)
- Kundurthy, P., Agol, E., Becker, A. C., et al. 2011, *ApJ*, 731, 123, doi: [10.1088/0004-637X/731/2/123](https://doi.org/10.1088/0004-637X/731/2/123)
- Kurucz, R. L. 1992, in *IAU Symposium*, Vol. 149, The Stellar Populations of Galaxies, ed. B. Barbuy & A. Renzini, 225
- Latham, D. W., Rowe, J. F., Quinn, S. N., et al. 2011, *The Astrophysical Journal*, 732, L24, doi: [10.1088/2041-8205/732/2/L24](https://doi.org/10.1088/2041-8205/732/2/L24)
- Li, J., Tenenbaum, P., Twicken, J. D., et al. 2019, *PASP*, 131, 024506, doi: [10.1088/1538-3873/aaf44d](https://doi.org/10.1088/1538-3873/aaf44d)
- Lightkurve Collaboration, Cardoso, J. V. d. M., Hedges, C., et al. 2018, *Lightkurve: Kepler and TESS time series analysis in Python*, *Astrophysics Source Code Library*. <http://ascl.net/1812.013>
- Lissauer, J. J., Ragozzine, D., Fabrycky, D. C., et al. 2011, *ApJS*, 197, 8, doi: [10.1088/0067-0049/197/1/8](https://doi.org/10.1088/0067-0049/197/1/8)
- Lissauer, J. J., Marcy, G. W., Rowe, J. F., et al. 2012, *ApJ*, 750, 112, doi: [10.1088/0004-637X/750/2/112](https://doi.org/10.1088/0004-637X/750/2/112)
- Lissauer, J. J., Marcy, G. W., Bryson, S. T., et al. 2014, *The Astrophysical Journal*, 784, 44, doi: [10.1088/0004-637x/784/1/44](https://doi.org/10.1088/0004-637x/784/1/44)

- Lithwick, Y., Xie, J., & Wu, Y. 2012, *ApJ*, 761, 122, doi: [10.1088/0004-637X/761/2/122](https://doi.org/10.1088/0004-637X/761/2/122)
- Lovis, C., & Fischer, D. 2010, *Radial Velocity Techniques for Exoplanets*, ed. S. Seager, 27–53
- Luger, R., Agol, E., Foreman-Mackey, D., et al. 2019, *AJ*, 157, 64, doi: [10.3847/1538-3881/aae8e5](https://doi.org/10.3847/1538-3881/aae8e5)
- Luque, R., Pallé, E., Kossakowski, D., et al. 2019, *Astronomy & Astrophysics*, 628, A39, doi: [10.1051/0004-6361/201935801](https://doi.org/10.1051/0004-6361/201935801)
- Mandel, K., & Agol, E. 2002, *The Astrophysical Journal*, 580, L171–L175, doi: [10.1086/345520](https://doi.org/10.1086/345520)
- McKinney, W. 2010, in *Proceedings of the 9th Python in Science Conference*, ed. S. van der Walt & J. Millman, 51–56
- Minkowski, R. L., & Abell, G. O. 1963, *The National Geographic Society-Palomar Observatory Sky Survey*, ed. K. A. Strand, 481
- Morton, T. D. 2012, *ApJ*, 761, 6, doi: [10.1088/0004-637X/761/1/6](https://doi.org/10.1088/0004-637X/761/1/6)
- . 2015, *isochrones: Stellar model grid package*. <http://ascl.net/1503.010>
- Motalebi, F., Udry, S., Gillon, M., et al. 2015, *Astronomy & Astrophysics*, 584, A72, doi: [10.1051/0004-6361/201526822](https://doi.org/10.1051/0004-6361/201526822)
- Moultaka, J., Ilovaisky, S., Prugniel, P., & Soubiran, C. 2004, in *SF2A-2004: Semaine de l’Astrophysique Française*, ed. F. Combes, D. Barret, T. Contini, F. Meynadier, & L. Pagani, 547
- Niraula, P., Redfield, S., Dai, F., et al. 2017, *The Astronomical Journal*, 154, 266, doi: [10.3847/1538-3881/aa957c](https://doi.org/10.3847/1538-3881/aa957c)
- Oliphant, T. 2006, *Guide to NumPy*
- Pecaut, M. J., & Mamajek, E. E. 2013, *ApJS*, 208, 9, doi: [10.1088/0067-0049/208/1/9](https://doi.org/10.1088/0067-0049/208/1/9)
- Pecaut, M. J., Mamajek, E. E., & Bubar, E. J. 2012, *ApJ*, 746, 154, doi: [10.1088/0004-637X/746/2/154](https://doi.org/10.1088/0004-637X/746/2/154)
- Perruchot, S., Kohler, D., Bouchy, F., et al. 2008, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 7014, *The SOPHIE spectrograph: design and technical key-points for high throughput and high stability*, 70140J, doi: [10.1117/12.787379](https://doi.org/10.1117/12.787379)
- Perryman, M., Hartman, J., Bakos, G. Á., & Lindegren, L. 2014, *The Astrophysical Journal*, 797, 14, doi: [10.1088/0004-637x/797/1/14](https://doi.org/10.1088/0004-637x/797/1/14)
- Petigura, E. A. 2005, PhD thesis, University of California, Berkeley
- Petigura, E. A., Crossfield, I. J. M., Isaacson, H., et al. 2017, *The Astronomical Journal*, 155, 21, doi: [10.3847/1538-3881/aa9b83](https://doi.org/10.3847/1538-3881/aa9b83)
- Price-Whelan, A. M., Sipőcz, B. M., Günther, H. M., et al. 2018, *AJ*, 156, 123, doi: [10.3847/1538-3881/aabc4f](https://doi.org/10.3847/1538-3881/aabc4f)
- Quinn, S. N., Becker, J. C., Rodriguez, J. E., et al. 2019, *AJ*, 158, 177, doi: [10.3847/1538-3881/ab3f2b](https://doi.org/10.3847/1538-3881/ab3f2b)
- Quirrenbach, A., Amado, P. J., Caballero, J. A., et al. 2014, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 9147, *CARMENES instrument overview*, 91471F, doi: [10.1117/12.2056453](https://doi.org/10.1117/12.2056453)
- Quirrenbach, A., Amado, P. J., Ribas, I., et al. 2018, in *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 10702, *Proc. SPIE*, 107020W, doi: [10.1117/12.2313689](https://doi.org/10.1117/12.2313689)
- Ragozzine, D., & Holman, M. J. 2010, *The Value of Systems with Multiple Transiting Planets*. <https://arxiv.org/abs/1006.3727>
- Reid, I. N., Brewer, C., Brucato, R. J., et al. 1991, *PASP*, 103, 661, doi: [10.1086/132866](https://doi.org/10.1086/132866)
- Rein, H., & Liu, S. F. 2012, *A&A*, 537, A128, doi: [10.1051/0004-6361/201118085](https://doi.org/10.1051/0004-6361/201118085)
- Rice, K., Malavolta, L., Mayo, A., et al. 2019, *MNRAS*, 484, 3731, doi: [10.1093/mnras/stz130](https://doi.org/10.1093/mnras/stz130)
- Ricker, G. R., Winn, J. N., Vanderspek, R., et al. 2014, *Journal of Astronomical Telescopes, Instruments, and Systems*, 1, 014003, doi: [10.1117/1.jatis.1.1.014003](https://doi.org/10.1117/1.jatis.1.1.014003)
- Rowe, J. F., Bryson, S. T., Marcy, G. W., et al. 2014, *The Astrophysical Journal*, 784, 45, doi: [10.1088/0004-637x/784/1/45](https://doi.org/10.1088/0004-637x/784/1/45)
- Salvatier, J., Wiecki, T. V., & Fonnesbeck, C. 2016, *PeerJ Computer Science*, 2, e55
- Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, *ApJ*, 500, 525, doi: [10.1086/305772](https://doi.org/10.1086/305772)
- Seader, S., Tenenbaum, P., Jenkins, J. M., & Burke, C. J. 2013, *ApJS*, 206, 25, doi: [10.1088/0067-0049/206/2/25](https://doi.org/10.1088/0067-0049/206/2/25)
- Seager, S. 2010, *Exoplanet Atmospheres: Physical Processes*, Princeton Series in Astrophysics (Princeton University Press). <https://books.google.com/books?id=XpaYJD7IE20C>
- Seager, S., & Mallen-Ornelas, G. 2003, *The Astrophysical Journal*, 585, 1038–1055, doi: [10.1086/346105](https://doi.org/10.1086/346105)
- Siverd, R. J., Brown, T. M., Hygelund, J., et al. 2016, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 9908, *NRES: the network of robotic Echelle spectrographs*, 99086X, doi: [10.1117/12.2233188](https://doi.org/10.1117/12.2233188)
- Siverd, R. J., Brown, T. M., Barnes, S., et al. 2018, in *Ground-based and Airborne Instrumentation for Astronomy VII*, ed. C. J. Evans, L. Simard, & H. Takami, Vol. 10702, *International Society for Optics and Photonics (SPIE)*, 1918 – 1929, doi: [10.1117/12.2312800](https://doi.org/10.1117/12.2312800)

- Spada, F., Demarque, P., Kim, Y. C., & Sills, A. 2013, *ApJ*, 776, 87, doi: [10.1088/0004-637X/776/2/87](https://doi.org/10.1088/0004-637X/776/2/87)
- Stassun, K. G., Collins, K. A., & Gaudi, B. S. 2017, *The Astronomical Journal*, 153, 136, doi: [10.3847/1538-3881/aa5df3](https://doi.org/10.3847/1538-3881/aa5df3)
- Stassun, K. G., Corsaro, E., Pepper, J. A., & Gaudi, B. S. 2018, *AJ*, 155, 22, doi: [10.3847/1538-3881/aa998a](https://doi.org/10.3847/1538-3881/aa998a)
- Stassun, K. G., & Torres, G. 2016, *The Astronomical Journal*, 152, 180, doi: [10.3847/0004-6256/152/6/180](https://doi.org/10.3847/0004-6256/152/6/180)
- Stassun, K. G., & Torres, G. 2018, *ApJ*, 862, 61, doi: [10.3847/1538-4357/aacafc](https://doi.org/10.3847/1538-4357/aacafc)
- Stassun, K. G., Oelkers, R. J., Pepper, J., et al. 2018, *The Astronomical Journal*, 156, 102, doi: [10.3847/1538-3881/aad050](https://doi.org/10.3847/1538-3881/aad050)
- Theano Development Team. 2016, arXiv e-prints, [abs/1605.02688](https://arxiv.org/abs/1605.02688). <http://arxiv.org/abs/1605.02688>
- Torres, G., Andersen, J., & Giménez, A. 2010, *A&A Rv*, 18, 67, doi: [10.1007/s00159-009-0025-1](https://doi.org/10.1007/s00159-009-0025-1)
- Twicken, J. D., Catanzarite, J. H., Clarke, B. D., et al. 2018, *Publications of the Astronomical Society of the Pacific*, 130, 064502, doi: [10.1088/1538-3873/aab694](https://doi.org/10.1088/1538-3873/aab694)
- Van Grootel, V., Gillon, M., Valencia, D., et al. 2014, *ApJ*, 786, 2, doi: [10.1088/0004-637X/786/1/2](https://doi.org/10.1088/0004-637X/786/1/2)
- van Leeuwen, F. 2007, *A&A*, 474, 653, doi: [10.1051/0004-6361:20078357](https://doi.org/10.1051/0004-6361:20078357)
- Vanderburg, A., Bieryla, A., Duev, D. A., et al. 2016, *The Astrophysical Journal*, 829, L9, doi: [10.3847/2041-8205/829/1/19](https://doi.org/10.3847/2041-8205/829/1/19)
- Virtanen, P., Gommers, R., Oliphant, T. E., et al. 2020, *Nature Methods*, doi: <https://doi.org/10.1038/s41592-019-0686-2>
- William H. Press, Saul A. Teukolsky, W. T. V. B. P. F. 1992, *Numerical Recipes: The Art of Scientific Computing* (Cambridge University Press)
- Wolfgang, A., Rogers, L. A., & Ford, E. B. 2016, *The Astrophysical Journal*, 825, 19, doi: [10.3847/0004-637x/825/1/19](https://doi.org/10.3847/0004-637x/825/1/19)
- Yi, S., Demarque, P., Kim, Y.-C., et al. 2001, *ApJS*, 136, 417, doi: [10.1086/321795](https://doi.org/10.1086/321795)

APPENDIX

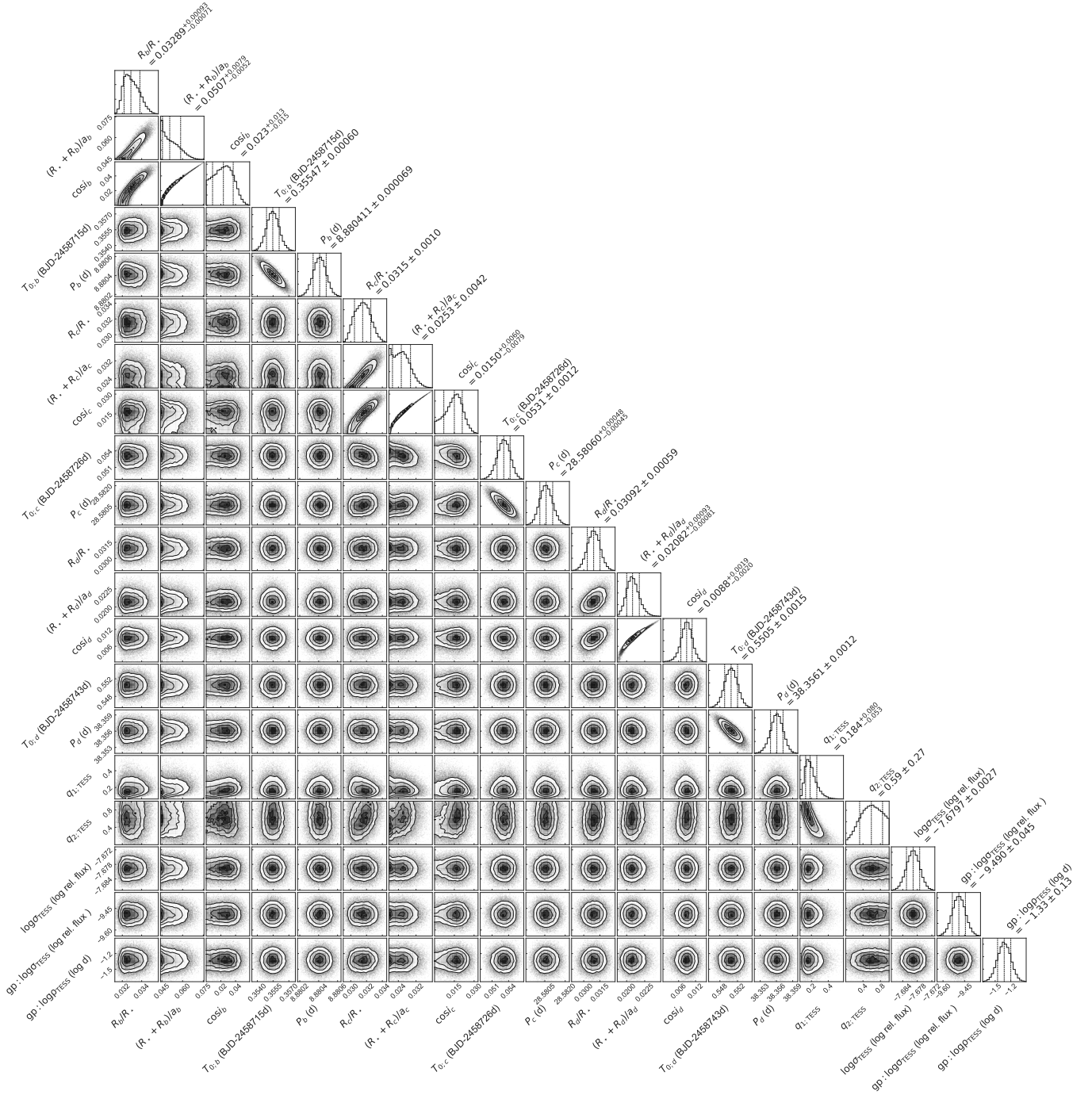


Figure 10. Posterior probability distributions for the `allesfitter` model parameters. The dashed lines show the 16th, 50th, and 84th percentiles.